

A catalogue of low-mass X-ray binaries*

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Received 30 October 2000 / Accepted 11 January 2001

Abstract. We present a catalogue of low-mass X-ray binaries. The catalogue is an updated version of the catalogue of van Paradijs (1995). This new catalogue contains 150 sources, 31 new low-mass X-ray binaries in addition to the 119 sources listed in van Paradijs’ catalogue. The aim of this catalogue is to help the reader gain easy access to the recent literature (up to about August 2000) on individual sources, and to provide of some basic information on the X-ray sources and their counterparts in other wavelength ranges (UV, optical, IR, and radio). In cases where there is some doubt about the low-mass nature of the X-ray binary, this is mentioned. In an appendix we list the Anomalous X-ray Pulsars (AXPs), which nowadays are no longer thought to be low-mass X-ray binaries.

Key words. X-ray: stars – (stars:) binaries: general – catalogs – pulsars: general

1. Introduction

An X-ray binary contains either a neutron star (NS) or a black hole (BH) accreting material from a companion star. Thus, cataclysmic variables (CVs) are not included among the X-ray binaries. X-ray binaries can be further divided into two different classes according to the mass of the companion star: high-mass X-ray binaries (HMXB) and low-mass X-ray binaries (LMXB). In a previous paper we presented an up-to-date catalogue of high-mass X-ray binaries (Liu et al. 2000).

The secondary of LMXB systems is a low-mass (in general $M \leq 1 M_{\odot}$) star, which transfers matter by Roche-lobe overflow. Among the low-mass companion stars we find white dwarfs, late-type main-sequence stars, two A-type stars and F-G-type sub-giants. The last-mentioned category of companion stars may well be the mass-transfer remnants of stars that originally were of intermediate mass ($M \sim 1.5$ to $4 M_{\odot}$) as has recently been suggested for Cygnus X-2 (cf. Podsiadlowski & Rappaport 2000; Tauris et al. 2000). The optical counterparts of LMXBs are intrinsically faint objects. The spectra of most of them show a few characteristic emission lines superposed on a rather

flat continuum. The optical continuum of LMXBs is dominated by the emission from an accretion disk around the compact star, which is predominantly the result of reprocessing of a fraction of the X-rays into optical photons in the disk. The contribution from the secondary is generally negligible. On occasion, however, the presence of the secondary can be discerned in the spectrum (or colors) of the LMXBs. This particularly is the case for systems with donors that are or started out as intermediate mass stars, such as Her X-1 and Cyg X-2. For a full understanding of LMXBs in various aspects one can refer to the book by Lewin et al. (1995).

The classification as LMXB is mainly based on the spectra obtained from an optical identification, and/or on the mass function from X-ray pulse arrival time measurements. If neither is available, a classification may be inferred based on the similarity of the X-ray properties to other identified systems. An unidentified system is classified as an LMXB containing a neutron star if one or more of the following properties are observed:

- type I X-ray bursts (which to date have only been seen from neutron stars in LMXBs);
- the 1–10 keV spectrum is soft with a characteristic temperature of 5–10 keV and/or;
- the orbital period is less than about 12 hr.

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* Table 1 is also available in electronically at the CDS via anonymous ftp (130.79.128.5) or via
<http://cdsweb.u-strasbg.fr/cgi-bin/qcat?J/A+A/368/1021>

In 1983 the number of known LMXBs was about 33 (van Paradijs 1983; Bradt & McClintock 1983). By the time of the previous catalogue the number of X-ray sources associated with low-mass stars had increased to 119 (van Paradijs 1995) (with three supersoft X-ray sources (SSSs) and two AXPs excluded (see Sect. 3)). In this paper we present a new catalogue, which includes the information extracted from the literature up to August 2000 and increases the number of known LMXBs to 150, including 31 newly discovered LMXBs as well as the 119 “old” ones listed in van Paradijs’ catalogue.

2. Description of the table

Table 1 lists the 150 LMXBs. The format of the table is similar to that of the previous one (van Paradijs 1995) and the well-known work of Bradt & McClintock (1983), of which the present catalogue is meant to be an update. In the table the sources are ordered according to right ascension; part of the (mainly numerical) information on a source is arranged in seven columns, below which for each source additional information is provided in the form of key words with reference numbers [in square brackets]. When a result is unreliable, a colon (:) or a question mark (?) will follow the adopted entry. The columns have been arranged as follows:

In Col. 1 the first line contains the source name, with rough information on its sky location according to the convention hhmm \pm ddd. Here hh and mm indicate the hours and minutes of right ascension, ddd the declination in units of 0.1 degree (in a small number of cases, the coordinates shown in the name are given with more, or fewer, digits). However, for a ROSAT source the name is always given in the form of hhmm.m \pm ddmm. The prefix J indicates a name based on J2000 coordinates. Otherwise, 1950 coordinates were used in the name. Alternative source names are given in the second line. The third line of Col. 1 lists survey catalogues and experiments in which the source was listed and detected, respectively. The following abbreviations have been used.

- A: Ariel V sky survey;
- AS: ASCA;
- B: BeppoSAX;
- C: Compton γ -ray Observatory;
- E: Einstein Observatory;
- Exo: Exosat;
- G: Ginga;
- Gr: Granat;
- H: HEAO A-1 sky survey;
- Ha: Hakuto;
- I: Indian X-ray Astronomy Experiment (IXAE);
- K: Kvant;
- M: MIT OSO-7 sky survey;
- OAO: Orbiting Astronomical Observatory;
- R: ROSAT;
- S: SAS 3;
- SL: Space Lab;

- T: Tenma;
- U: Uhuru sky survey;
- V: Vela-5 and -6 satellites;
- X: Rossi XTE.

In the first line of Col. 2, the source types are indicated with a letter code, as follows:

- A: atoll source (18);
- B: X-ray burst source (63);
- D: “dipping” low-mass X-ray binary (11);
- G: globular-cluster X-ray source (13);
- P: X-ray pulsar (5);
- T: transient X-ray source (76);
- U: ultra-soft X-ray spectrum. These sources include black-hole candidates; some “extreme ultra-soft” (EUS) sources may be white dwarfs on whose surface steady nuclear burning takes place.
- Z: Z-type (7).

In the third line of Col. 2, we provide some information on the type of observation from which the source position has been derived. The following abbreviations have been used: o, optical; x, X-ray; r, radio; IR, infrared. A reference on the source position is given below the columnar information under “Pos.”. In addition, we give an indication of the accuracy of this position, in the form of equivalent (90 percent confidence level) error radii, but in several cases this can only be considered an approximation (e.g. when the error box is not circular). When no accuracy is quoted, it is about one arcsecond or better.

Column 3 contains in the first two lines the right ascension (RA) and declination (DEC) of the source for epoch 1950 for the usual name, and for epoch 2000 for the sources with the names of J2000 coordinates. RA is given as hhmmss.s to an accuracy of 0.1 s, DEC is given in $^{\circ} \text{ } ' \text{ } ''$, to an accuracy of $1''$. The third line gives the galactic longitude and latitude to an accuracy of 0.1° (except for sources close to the galactic center (GC), where these coordinates are given to 0.01°).

The first and second lines of Col. 4 give names of an optical counterpart. The third line contains a reference to a finding chart. An asterisk followed by a number or letter refers to star numbers used in the finding chart; “star” refers to a star in the finding chart that has not been assigned a number or letter. Many optical counterparts have been indicated with a variable-star name, as given in the *General Catalogue of Variable Stars* and in recent name lists of variable stars as published regularly in the *IAU Information Bulletin on Variable Stars*, or a number in a well-known catalogue (e.g. HD, SAO). For X-ray sources in globular clusters, the cluster name is here given, in addition to the name of a stellar optical counterpart.

The fifth column contains some photometric information on the optical counterpart. In the first line, the apparent visual magnitude, V , and the color indices $B - V$, and $U - B$, are listed. The second line contains the estimate of the interstellar reddening, E_{B-V} .

Table 1. Low-mass X-ray binaries

Table 1. continued

In Col. 6, the average X-ray flux, or the range of observed X-ray fluxes (2–10 keV, unless otherwise indicated), is given, in units of

$$\begin{aligned} 1 \mu\text{Jy} &= 10^{-29} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1} \\ &= 2.4 \cdot 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ keV}^{-1}. \end{aligned}$$

The first line in Col. 7 gives the orbital period in hours. The second line contains for X-ray pulsars the pulse period, in seconds.

3. Remarks on some individual objects

We wish to emphasize here that some sources listed in this catalogue are still uncertain. They need to be regarded with caution in view of the further work necessary. Some sources are tentatively classified as low-mass X-ray binaries due to their X-ray burst feature and/or the similarity of the X-ray properties to other identified systems. No counterpart at other bands has been found.

Finally, we would like to make some remarks on several individual sources. The three luminous supersoft X-ray sources, RX J0527.8–6954, 1E 0543.8–6823 (CAL 83) and 1E 0547.6–7109 (CAL 87), are excluded from the new catalogue. The now generally accepted model for the luminous SSSs is that, with few exceptions, these sources are accreting white dwarfs (WDs) in binaries, which are burning hydrogen in their envelopes in a steady or intermittent way (van den Heuvel et al. 1992; Kahabka & van den Heuvel 1997).

There is a new class of pulsating objects with spin periods in the 6–12 s range, which are different from “normal” binary X-ray pulsars (Mereghetti & Stella 1995). These objects are referred to as braking X-ray pulsars (BXP) (Mereghetti & Stella 1995) or anomalous X-ray pulsars (van Paradijs et al. 1995). These sources include 4U 0142+61, 1E 2259+586, 1E 1048.1–5937, RX J1838.4–0301 (but see Mereghetti et al. 1997; Song et al. 2000 for RX J1838.4–0301) and the recently discovered sources 1E 1841–045 (Vasisht & Gotthelf 1997), 1RXS J170849.0–400910 (Sugizaki et al. 1997), AX J1845.0–0358 (Torii et al. 1998), and probably RX J0720.4–3125 (Haberl et al. 1997). 4U 0142+61 and 1E 2259+586 have already been listed in the previous LMXB catalogue, while 1E 1048.1–5937 was listed in the HMXB catalogue. Although it has been suggested that AXPs may consist of neutron stars accreting from a very low-mass companion (Mereghetti & Stella 1995) or are isolated neutron stars accreting from circumstellar debris (e.g. Corbet et al. 1995; van Paradijs et al. 1995), recently, evidence has been mounting that they are isolated

neutron stars with very high magnetic fields, i.e., magnetars related to the soft gamma-ray repeaters (SGRs) (Kouveliotou et al. 1998, 1999; Hurley et al. 1999), as has been suggested by Thompson & Duncan (1996). In view of these developments we have not included the AXPs in this LMXB catalogue, but in an Appendix we present a table of the suspected AXPs.

Acknowledgements. We are grateful to Robert Fender and Michiel van der Klis for carefully reading the manuscript and for useful comments. QZL acknowledges the financial support from KC Wong fellowship of Chinese Academy of Sciences. This work is partially supported by the Netherlands Organization for Scientific Research (NWO) through Spinoza Grant 08–0 to E. P. J. van den Heuvel and by the National Project for Fundamental Research by the Ministry of Science and Technology of China (973 Project) through Grant G1999075405.

Noted added in proof. Rudy Wijnands informed us that we had missed two new sources SAX J1753.5–2349 and SAX J1806.5–2215 (in't Zand et al. 1998, *axrs. symp.* 228), while the sources SAX J0840.7+2248 (IAUC 6892), SAX J1719.6–4254 (IAUC 7263), XTE J1739–285 (IAUC 7300), XTE J1743–363 (IAUC 7210), and SAX J1818.6–1703 (IAUC 6840) are also likely LMXBs.

Appendix: Anomalous X-ray pulsars

There is growing evidence that anomalous X-ray pulsars are magnetars, a type of objects with dipolar magnetic fields much stronger than the critical magnetic field. Their common properties are:

- (1) their spin periods are distributed in a narrow range (5–12 s), contrary to those of high-mass X-ray binary pulsars, which cover a much wider range;
- (2) they have very soft X-ray spectra with a photon index about 3 or larger;
- (3) their X-ray luminosity is relatively low (10^{35} – 10^{36} erg s $^{-1}$), compared to that of HMXB pulsars;
- (4) their X-ray flux shows little variability on timescales from months to years (they are not transient systems);
- (5) they have a relatively stable spin period evolution, with long intervals of nearly constant spin-down;
- (6) they are located in the galactic plane with the galactic latitude $|b| \leq 1^\circ$;
- (7) their optical counterparts, if present, are very faint of implying that at least they cannot have a massive companion; and
- (8) a few of them are possibly associated with supernova remnants (SNRs).

In this appendix we present a simple table of these anomalous X-ray pulsars (Table 2). Further identification is needed for some AXP candidates, although we have

Table 2. Anomalous X-ray pulsars

Pulsars	l^{II} , b^{II}	P (s)	\dot{P} ($10^{-11} \text{ s s}^{-1}$)	B_{P} (10^{14} G)	kT (keV)	L_{X} ($10^{35} \text{ erg s}^{-1}$)	d (kpc)	z (pc)
4U 0142+615	129.4, -0.4	8.69	~ 0.2	2.6	0.39	1.4	1.5	-11
Pos.: [1419]; pulsation: [622]; finding chart: [1419]; X-ray obs.: [521, 623, 1014, 1419, 1443].								
RX J0720.4-3125	244.1, -8.1	8.39	0.26	~ 1	0.08	0.002	0.3	-43
Pos.: [486]; pulsation: [486]; finding chart: [486]; X-ray obs.: [486, 1398]; magnetic field: [535]; optical obs.: [706, 922].								
1E 1048.1-5937	288.3, -0.5	6.44	1.5-4	>6.3	0.64	5.6	10.6	+93
Pos.: [1136]; pulsation: [1136]; finding chart: [885]; X-ray obs.: [1136, 977]; limits on the orbital parameters: [888].								
RX J170849-4009	346.5, +0.02	11.0	2	9.5	0.40	10	10	+3
Pos.: [1227]; pulsation: [1227]; finding chart: [1227]; X-ray obs.: [1227].								
RX J1838.4-0301	28.8, +1.5	5.45	—	—	—	0.6	5.2	+135
Pos.: [1132]; pulsation: [1132]; pulsation not detected: [1197]; finding chart: [886]; X-ray obs.: [886, 1197].								
1E 1841-045	27.7, +0.26	11.76	4.7	14	0.55	3.5	7.0	+32
Pos.: [1362]; pulsation: [1362]; X-ray obs.: [441, 1362]; associated with young SNR Kes 73: [1362].								
AX J1845.0-0300	29.52, +0.07	6.97	—	—	0.64	0.91	8.5	+10
Pos.: [1294]; pulsation: [1294]; X-ray obs.: [1294]; associated with SNR (?): [392].								
1E 2259+586	109.1, -1.0	6.98	~ 0.05	1.2	0.41	1.7	6.2	-108
Pos.: [291]; pulsation: [292]; finding chart: [291]; X-ray obs.: [58, 103, 663]; spin period var.: [57, 58, 240]; X-ray glitcher: [1316]; pulse period history: [292, 699, 700, 943]; associated with SNR G109-1 (?): [457, 587]; no optical/IR counterpart: [291, 671]; X-ray cyclotron line: [624]; deep radio search: [223]; limits on the orbital parameters: [888].								

included them in the table. Of these, AX J1845-0258 and RX J1838.4-0301 have no detected meaningful period derivatives due to the limited time span of the data, while the enigmatic source RX J072.4-3125 has a period of 8.39 s, but its X-ray luminosity is too low for an AXP.

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