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A Digital Correlator Upgrade for the Arcminute MicroKelvin Imager

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ABSTRACT

The Arcminute Microkelvin Imager (AMI) telescopes located at the Mullard Radio Astronomy Observatory near Cambridge have been significantly enhanced by the implementation of a new digital correlator with 1.2 MHz spectral resolution. This system has replaced a 750-MHz resolution analogue lag-based correlator, and was designed to mitigate the effects of radio frequency interference, particularly from geostationary satellites that contaminate observations at low declinations. The upgraded instrument consists of 18 ROACH2 Field Programmable Gate Array platforms used to implement a pair of real-time FX correlators – one for each of AMI’s two arrays. The new system separates the down-converted RF baseband signal from each AMI receiver into two 2.3 GHz-wide sub-bands which are each digitized at 5-Gsps with 8 bits of precision. These digital data streams are filtered into 2048 frequency channels and cross-correlated using FPGA hardware, with a commercial 10 Gb Ethernet switch providing high-speed data interconnect. Images formed using data from the new digital correlator show over an order of magnitude improvement in dynamic range over the previous system. The ability to observe at low declinations has also been significantly improved.

Key words: instrumentation: interferometers – techniques: interferometric – telescopes

1 INTRODUCTION

The Arcminute Microkelvin Imager (AMI) instrument consists of two interferometric radio telescopes located at the Mullard Radio Astronomy Observatory near Cambridge. AMI, which operates in the 12–18 GHz frequency band, was designed primarily for the study of galaxy clusters by observing secondary anisotropies in the cosmic microwave background (CMB) arising from the Sunyaev-Zel’dovich (SZ) effect (Sunyaev & Zeldovich 1970).

AMI is made up of two arrays: the Small Array (SA), and the Large Array (LA). The SA telescope comprises ten 3.7 m diameter paraboloid dishes in a compact configuration (Figure 1) and is designed for observing structures on arcminute scales. The LA, which

was created by reconfiguring the eight 12.8 m dishes of the Ryle Telescope (Figure 2) has an angular resolution of 0.5 arcminutes and has approximately ten times the flux-density sensitivity of the SA. The LA, observing concurrently with the SA, is used to measure the intensities of contaminating small-diameter radio sources.

The original AMI correlator (Zwart et al. 2008) was of an analogue "XF" design, whereby the antenna signals were cross-correlated over a range of delays and then Fourier transformed into spectral channels. The original correlator recorded visibilities in eight frequency channels over AMI’s 12–18 GHz band, providing limited capability for recognising and removing interfering signals as well as limited ability to mitigate chromatic aberration out to the edge of field of view. As a result, the sensitivity of the instrument was significantly reduced, particularly at low declinations, where interference from geostationary satellites could result in up to 90% of the data being unusable. The XF system had a performance that

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Figure 1. The SA telescope is located inside a metal enclosure with a metallic ground plane which acts as a reflector and also minimises radio frequency interference (RFI).

was also limited by the path-length inaccuracies of analogue delay compensation and analogue correlation, which resulted in direction-dependent systematic errors and poor dynamic range.

Starting in 2014, a project was undertaken to build a digital correlator for the telescope based on the second-generation Reconfigurable Open Architecture Computing Hardware board (ROACH2¹), developed by the Collaboration for Astronomy Signal Processing and Electronics Research (CASPER²) and used extensively in radio-astronomy digital signal processing (DSP) applications (Hickish et al. 2016). ROACH2 is a powerful real-time signal processing platform based around a Xilinx Virtex 6 Field Programmable Gate Array (FPGA) which provides up to 80 Gbps of I/O, and may be interfaced with a variety of analogue to digital converters (ADCs). The primary goal of the project was to equip AMI with a digital correlator providing over one hundred times the frequency resolution of the original system with superior inter-channel rejection, allowing better frequency-confinement of interfering signals as well as overcoming some of the limitations of analogue lag correlator. The wide bandwidth of the AMI receiver required an optimised FPGA correlator design, capable of operating at the relatively high clock speed of 312.5 MHz, to be developed. This correlator, which fits into a single 19-inch rack, comprises ten (SA) or eight (LA) ROACH2 boards which digitise, channelise and then cross-correlate analogue inputs from the antennas in the AMI arrays. Communication between boards in each AMI array is achieved using an industry-standard 10 Gb Ethernet (10 GbE) network, with an off-the-shelf switch providing interconnect. This so-called "Packetized Correlator" architecture was pioneered by the CASPER group (Parsons et al. 2008) and has been well proven at multiple radio-telescopes (Kocz et al. 2015; Foley et al. 2016; Nita et al. 2016; Primiani et al. 2016).

In this paper we summarize the design of, and first commissioning results from the new AMI correlator. In Section 2 we describe the architecture of the system deployed, which comprises an overhauled analogue front-end and a new digital back-end. In Section 3 details of the implementation of the digital back-end are presented. The instrument's control system and data reduction pipeline are outlined in Sections 4 and 5, respectively, with commissioning results from the new system given in Section 6.

¹ https://casper.berkeley.edu/wiki/ROACH-2_Revision_2

² <https://casper.berkeley.edu>



Figure 2. The LA telescope is located on the North-East site approximately 1.5 km away from the SA instrument. Figure credit: Nigel Blake

Table 1. The AMI frequency bands.

	RF (GHz)	IF (GHz)	Baseband (GHz)
Low Band	15.6–17.9	6.1–8.4	0.1–2.4
High Band	13.1–15.4	8.6–10.9	0.1–2.4

2 ARCHITECTURE OVERVIEW

2.1 Analogue Front-End

The AMI front-end system consists of a feed assembly, first down-converting mixer and second mixer stage as shown in Figure 3. Each feed assembly contains the horn antenna, cryostat, radio-frequency (RF) amplifiers, noise injection, bias box and the cryogenic support system. Each cryostat has three temperature zones separated by thermal shielding; 15–20 K, 50–70 K and 300 K. There are two RF amplifiers: the first is mounted on the 20 K coldhead of the cryostat, and the second is at 300 K. Noise at a constant level may be switched into the front-end via a noise source which is outside the cryostat on the SA, and inside the cryostat at 300 K on the LA. A detailed description of this system is given in Zwart et al. (2008).

The RF (12–18 GHz) signals from the antenna feeds are down-converted to an intermediate-frequency (IF) band of 6–12 GHz using a 24 GHz local oscillator (LO). Phase-switching is provided by modulating this LO signal by Walsh functions which are generated by the digital correlator. The IF is fed from the antenna hub to a correlator room where an Automatic Gain Control (AGC) unit is used to maintain a constant power level prior to further down-mixing.

A second mixer assembly splits the down-converted IF output into two 2.3 GHz wide sub-bands by first filtering the low and high bands and then mixing with an 8.5 GHz common local oscillator. Custom-manufactured low pass filters are used to further limit the bandpass of the baseband signal prior to being fed into ADCs. The terms "low band" and "high band" used here to describe the two AMI correlator basebands refer to the frequencies of the mixer bandpass filters, not their respective RF input frequencies as shown in Table 1. The second mixer assembly for the new system provides approximately 5 GHz of the original 6 GHz AMI frequency channel. The improvements in dynamic range are evident even with the loss in sensitivity due to processing less channel bandwidth. A picture of the new mixer assembly is shown in Figure 4.

2.2 Digital Back-End

The AMI digital correlator is implemented using the standard "FX" architecture favoured by most modern correlators. In this architecture analogue signals from each antenna are first digitised before

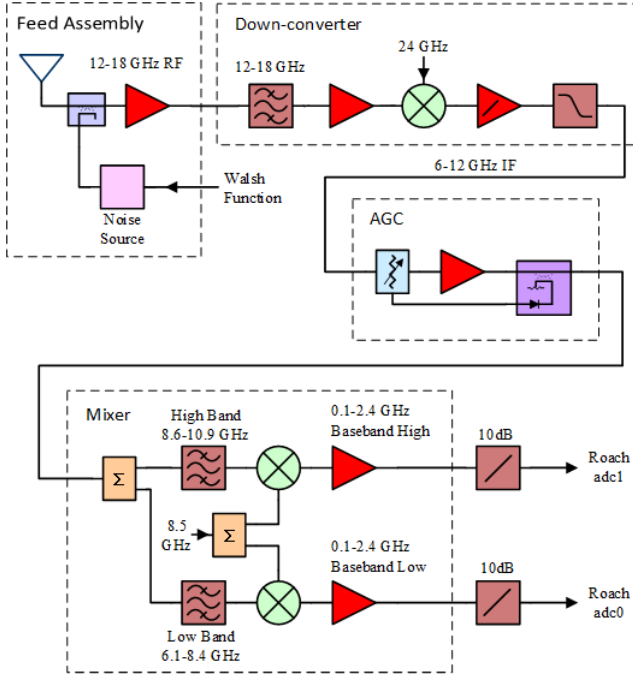


Figure 3. The upgraded AMI RF and IF system diagram. Isolators and attenuators used in the chain are not shown.

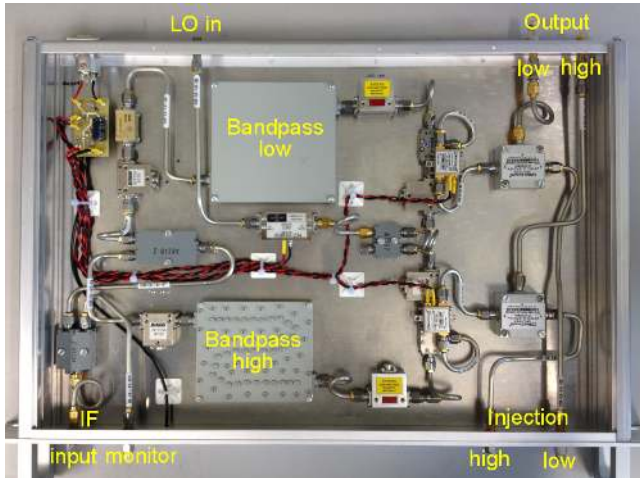


Figure 4. The upgraded AMI mixer assembly in a 1U enclosure for each antenna.

being split into spectral channels using an FFT-based algorithm (the "F" processing stage). Spectra from pairs of antennas are then multiplied and time-averaged on a per-frequency-channel basis (the "X" processing stage). The F-stage processing of such a correlator can be trivially parallelised over the multiple antennas in an array, whilst the X-stage computation is parallel over multiple frequency channels. Interconnect exists between the F and X processors to facilitate a data transpose—often referred to as a *corner-turn*—which allows data to be aggregated appropriately so as to allow parallel processing.

The AMI correlator uses FPGAs to perform the F and X processing, owing to their support for high input and output data rates and ease of interfacing to high-speed ADCs. Interconnect

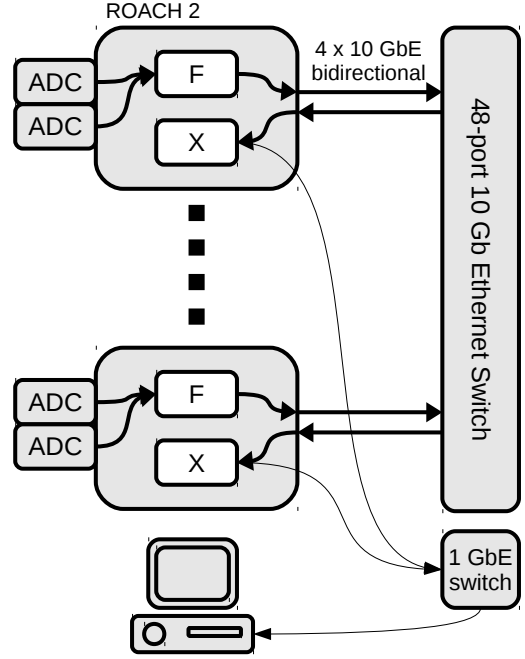


Figure 5. The top-level correlator architecture. Digitisation is performed by a pair of 5 Gbps ADC cards interfaced to a ROACH2. The ROACH2 firmware implements both channelisation and cross-multiplication functionality, with a 10 GbE switch providing interconnect between the two phases of processing. Averaged correlation results are output over a separate 1 Gb Ethernet network.

is provided by commercial off-the-shelf (COTS) 10 Gb/s Ethernet (10 GbE) switches, which offer cost-effective, flexible and reliable performance. Such switches also support useful functionality such as multicast (point to multi-point data transmission) which may be used in future correlator operating modes. This “packetised correlator” architecture has been well-developed by the CASPER collaboration, which provides and maintains software and firmware libraries to facilitate rapid deployment of digital astronomical systems.

Though not required by the packetised architecture, the AMI correlator design uses the same physical processors for its F and X stages (Figure 5). This allows the complete system to be efficiently implemented with minimal hardware. Once correlation matrices have been computed and averaged for an appropriate time window, these are sent over a 1 Gb Ethernet network where the results are aggregated by a standard x86 GNU/Linux server.

2.2.1 Digital Hardware

The digital processing platform chosen for the AMI correlator is the mature and ubiquitous CASPER-designed ROACH2 board³. This is a general-purpose FPGA platform, built around a Xilinx Virtex 6 (XC6VSX475T-1FFG1759C) FPGA, supported by four 72 Mb QDR memory chips, 1 GB of DRAM, and up to eight 10 GbE interfaces via a pair of SFP+ mezzanine cards. The ROACH2 also provides two 40-pair LVDS interfaces via Z-DOK connectors, which can be used to interface the ROACH2’s FPGA to a variety of CASPER-supported ADC cards.

For the AMI correlator, both Z-DOKs are populated with

³ <https://casper.berkeley.edu/wiki/ROACH2>

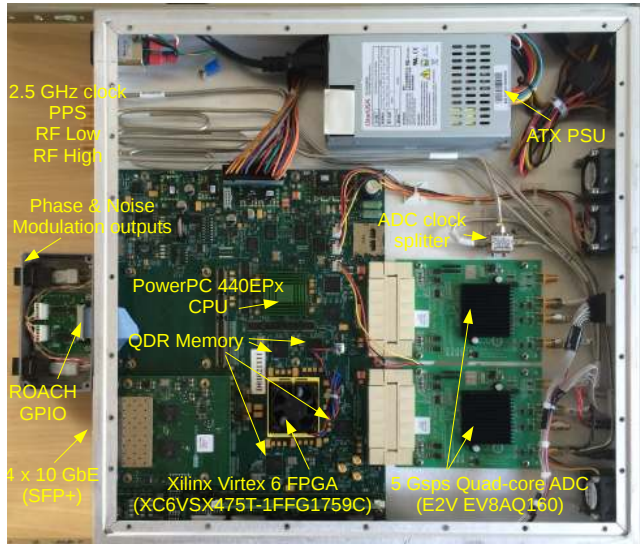


Figure 6. The CASPER ROACH2 platform on which the AMI digital correlator is implemented. Here the board is shown with two 5 Gbps ADC daughter cards, a quad-SFP+ mezzanine card, and a Walsh switch interface box.

CASPER ADC1x5000⁴ cards. Each of these hosts an e2v EV8AQ160: a quad-core digital sampler capable of sampling a single RF input at up to 5 Gbps. The ADC1x5000 was designed by the Academia Sinica Institute of Astronomy and Astrophysics (ASIAA) and has been extensively characterised by Patel et al. (2014) as part of the development of a new wideband correlator for the Submillimetre Array (SMA, Primiani et al. 2016).

Each ROACH2 in the AMI system is fed with two timing signals derived from a COTS Trimble Thunderbolt E GPS-disciplined oscillator, which provides a pulse-per-second (PPS) reference and 10 MHz frequency standard. The PPS is distributed to each ROACH2 in the system via a 16-way buffered splitter and is used to synchronise and timestamp the data outputs from each board. The 10 MHz reference is used to derive a 2500 MHz clock, which is amplified and split so that it may drive the 5000 Msp ADC samplers synchronously.

The complete ROACH2 processing node is shown in Figure 6. The node interfaces comprise:

- (i) An SMA input for a 2.5 GHz clock (used to derive the 5 Gbps ADC sampling rate).
- (ii) An SMA input for a TTL Pulse-per-second (PPS) time reference.
- (iii) SMA inputs for the high and low baseband inputs from the analogue front-end of a single AMI antenna.
- (iv) Level-shifting circuitry to allow interfacing of the ROACH2 FPGA with the AMI phase and noise modulation infrastructure (see Section 3.2).
- (v) A 1000BASE-T Ethernet interface to the FPGA, for data output.
- (vi) A 1000BASE-T Ethernet interface to ROACH2's on-board CPU, which is used for controlling and monitoring the board.
- (vii) Four SFP+ connectors, each providing a 10 GbE interface.

A key part of the correlator hardware infrastructure is the

⁴ <https://casper.berkeley.edu/wiki/ADC1x5000-8>

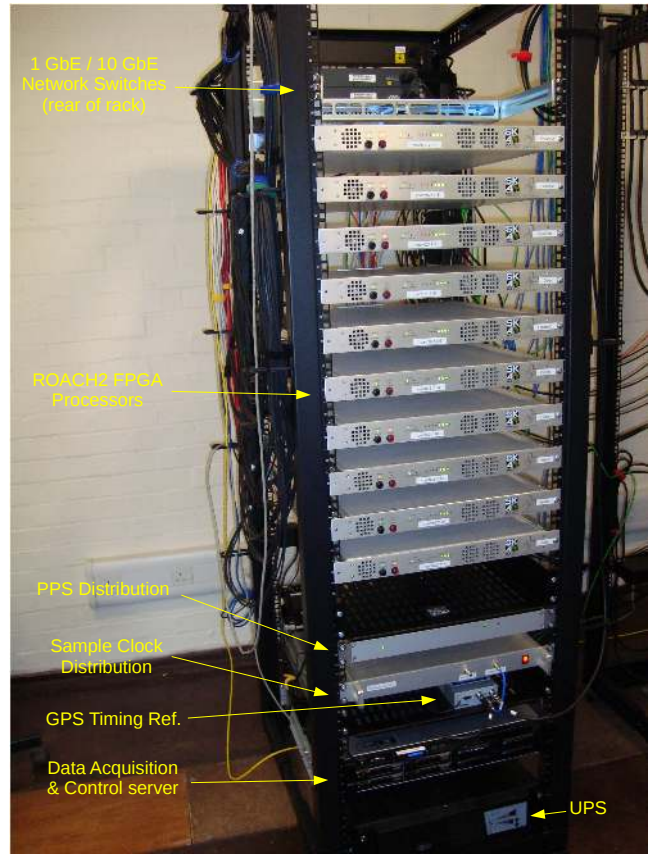


Figure 7. The digital correlator rack for the 10-antenna AMI Small Array.

10 GbE switch facilitating interconnection between the 10 (8) boards in the SA (LA) system. This is a Mellanox SX1012, 12-port, 40 Gb Ethernet switch, capable of operating as a 48-port 10 GbE switch using interconnecting cables to connect each of the switch's QSFP+ ports to four independent SFP+ interfaces.

The complete rack of digital equipment for the SA correlator is shown in Figure 7.

3 CORRELATOR FIRMWARE

3.1 Analogue-to-Digital Converters

The CASPER collaboration (and in particular CASPER members in the Submillimetre Array correlator group) provide open-source interface firmware⁵ and software⁶ to stream data from the e2v EV8AQ160 ADC into the ROACH2 board via its Z-DOK connector. This interface expects data to be presented to the FPGA at a quarter of the ADC sample rate, over four parallel 8-bit buses. At a maximum sampling rate of 5 Gbps, this results in 32 parallel data lines, each running at 1.25 Gbps. Once captured by the FPGA, these parallel samples are demultiplexed by a further factor of four, so that on every FPGA clock cycle, 16 samples are processed in parallel. In the AMI design the signal processing pipeline is clocked synchronously with the ADC, at a rate of 312.5 MHz. At the time

⁵ https://github.com/casper-astro/mlib_devel

⁶ https://github.com/sma-wideband/adc_tests

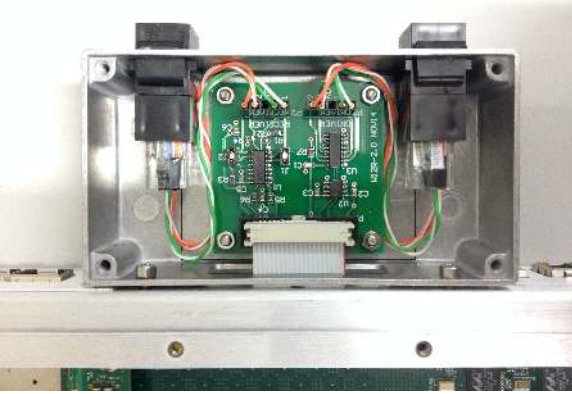


Figure 8. An interface board allowing driving of AMI’s 5 V differential phase and noise switch control signals from the 1.5 V ROACH GPIO. The interface also provides two inputs for optional use in the future.

of design, the interface provided by CASPER did not function at the full 5 Gbps rate of which the ADCs are capable. As part of this work, the interface has been modified to increase performance and released back to the CASPER community.

3.2 Front-end noise injection & phase switching control

In the original AMI analogue correlator, phase and noise modulation functions were generated by a bank of look-up tables driven by a 65536 Hz counter. In order to simplify the new digital system, it was decided that the correlator FPGAs should be responsible for generation of the switching signals. In this way the switching frequencies are synchronous with the ADC sampling, and can easily be chosen such that an integral number of switching periods occur in each channelisation window and correlator integration period.

In order that the FPGAs are able to drive the existing front-end switching infrastructure, simple interface boards (Figure 8) were constructed to convert a pair of 1.5 V ROACH2 general-purpose IO (GPIO) outputs to drive 5 V differential signals over Category 5 UTP cable. Though unused in the AMI correlator, the interface boards are also capable of converting a pair of differential signals to single-ended 1.5 V inputs, which may be used to drive ROACH2 GPIO pins.

3.3 F-Engine

The bulk of the new AMI correlator design comprises the “F-Engine” processing pipeline; the channelisation of pairs of 5 Gbps data streams into 2048 critically-sampled subbands, which are output over 10 GbE as User Datagram Protocol (UDP) data streams (Figure 9). In order to keep pace with the 5 Gbps ADCs used by the correlator, the AMI F-Engine pipeline, which runs at 312.5 MHz, processes 16 ADC samples on every FPGA clock tick. Stages of the pipeline are:

Phase-demodulation Immediately after samples are captured into the FPGA, the phase modulation applied in the antenna’s first mixer stage is removed by a simple multiplication of ADC samples by ± 1 . The modulation/demodulation pattern used is unique to each antenna, and stored in the firmware in a runtime-programmable look up table. A programmable delay between the modulation GPIO output (Section 3.2) and the internal demodulation signal allows

compensation for delays associated with cable lengths of the control and RF signals.

Coarse delay After demodulation, each antenna’s digital data stream may be delayed by up to 16,384 ADC samples. This delay allows for compensation of geometric delays in the array (the largest baseline in the LA is approximately 110 m) and RF cabling. The desired delays are calculated by the telescope control computer based on the current pointing of the array, and any pre-computed delay calibration solutions. The frequency resolution of the AMI digital correlator is 1.22 MHz, giving an inverse-channel-bandwidth of around 800 ns. This inverse-bandwidth sets the accuracy with which variable delays must be applied to avoid losing coherence of the antenna signals; for AMI, it is sufficient to update coarse delays on \sim second timescales. For simplicity of data analysis, the correlator control software ensures that delay updates are applied synchronously with new visibility accumulations, such that no coarse delay changes occur mid-integration.

Polyphase Filterbank The largest component of the F-Engine processing pipeline is a polyphase filterbank (originally proposed by Bellanger et al. (1976), see Harris & Haines (2011); Price (2016) for relevant overviews), which breaks the 5 Gbps data streams into 2048 critically sampled subbands, each of bandwidth 1.22 MHz. The CASPER group provide parameterised finite impulse response (FIR) filter and fast Fourier transform (FFT) libraries for implementing polyphase filterbanks – these have been utilised here, after various optimisations were applied to reduce FPGA resource utilisation and to maximise clock frequency.

Signals within the channeliser are processed with 18-bits of precision in both channeliser coefficients and data path, resulting in output subbands (spectra) which are complex-valued, with each of the real and imaginary components represented as signed 18-bit numbers.

Autocorrelation sub-system After channelisation, the power of each spectrum is computed and fed into a vector accumulator. This provides (without requiring the downstream correlator infrastructure) an averaged power-spectrum for each antenna signal. The autocorrelation subsystem can also demodulate by the Walsh pattern used to drive the noise injection at each of the antenna feeds, providing a measurement of the system-temperature of each antenna. Since the system temperature of antennas varies with weather, constant measurements of these “rain-gauge” values are performed and stored with final data products.

Requantization In order to reduce FPGA output bandwidth, data samples are rounded to 4-bit values – real and imaginary parts in the range $(-7, +7)$ – with saturation logic prior to being streamed over Ethernet. In order to use the available 4-bits of range most efficiently, before quantisation each of the 2048 subbands associated with an ADC data stream are equalised by multiplying by per-antenna, per-subband coefficients. These coefficients are runtime-programmable, and can be calculated either by examining data from the autocorrelation sub-system or the final correlator data products.

Data buffering and transmission The CASPER cross-correlation module used in the correlator (Parsons et al. 2008; Hickish 2014) expects data to be presented in windows of 1024 samples from each frequency channel in turn. To achieve this, 1024 spectra, each of 2048 channels, must be re-ordered in memory to gather similar channels into contiguous blocks. This reordering operation requires

a buffer of \sim MB size. Such memory is not available on the FPGA chip itself, so external 72 Mb Quad Data Rate (QDR) memory chips (of which four are provided by the ROACH2 platform) are used for this reorder operation.

Once data are suitably ordered, they are transmitted as streams of UDP packets over an Ethernet network. The destination address of each packet is determined by the frequency channel associated with the data in that packet’s payload, with the end result that each ROACH2 in the network receives data associated with all antennas in the array, but only certain frequency channels.

3.4 Data Interconnect

After 4-bit requantization, the data output rate from each dual-band F-engine is 40 Gbps. These data are transmitted as a stream of UDP packets, with each packet’s payload comprising 1024 time samples from a single frequency channel. When transmitted over Ethernet as a stream of UDP packets, transmission protocol overhead – Ethernet, Internet Protocol, UDP and application headers – must also be accommodated. In order to transmit the stream over four 10 GbE links the AMI system takes advantage of the fact that F- and X-processors are located on the same physical hardware. This allows frequency channels with the same source and destination ROACH2 to be routed internally, bypassing the Ethernet interconnect. Such routing reduces the necessary throughput of the interconnection by 10% (12.5%) for the Small (Large) AMI array. Including transmission protocol overhead, the aggregate output data rate from each 10 GbE interface is 9.4 (9.1) Gbps for the Small (Large) array.

3.5 X-Engine

The second stage in the correlator is the “X-Engine”, which is responsible for performing a per-frequency-channel cross-multiplication of data from different antenna pairs. The X-Engine pipeline is shown in Figure 10, and its components comprise:

Input buffering Data are received from the Ethernet network on the same four SFP+ interfaces used for transmission, as well as an internal routing path. The order that packets from different antenna sources are received from the network is not known in advance – a circular buffer is used to collect, and appropriately order, data from different antennas prior to cross-multiplication. Once all antenna packets have been received for a given frequency channel, n , which is indicated by the arrival of a packet of channel $n + 2$, a window of data is streamed into a cross-multiplication engine.

Cross-Multiplication The cross-multiplication engine is responsible for taking 1024 samples from a single frequency channel from all antennas, and delivering a visibility matrix integrated over these samples. The correlation engine used in AMI is based on a “windowed X-Engine” design by Parsons et al. (2008) which is now maintained by the CASPER community. The AMI version of this module has been written in Verilog (rather than the CASPER standard of Xilinx System Generator and MATLAB Simulink) making it more portable and maintainable than the original design. It is also significantly optimised to minimise FPGA resource utilisation and maximise performance.

Firstly, AMI’s X-engine has a parameterisable input bandwidth, which can be any multiple of the FPGA clock speed. This allows a reduction in control-logic versus multiple instances of a fixed input-bandwidth engine. Secondly, the cross-multiplication

and accumulation cores at the heart of the X-engine have been optimised for 4+4-bit complex inputs, allowing a 75% reduction in multiplier use. This is achieved by offsetting the 4-bit correlator inputs to unsigned values, and appropriately packing pairs of them in 18-bit representations such that four 4-bit multiplications may be computed in a single 18×18 -bit operation. A full description of this implementation can be found in Hickish (2014), which builds on the 4-bit packing speed-ups of de Souza et al. (2007) and is similar to more recent work targeting cross-multiplication of astronomical signals on Graphics Processing Units (GPUs) by Klages et al. (2015). As a result, the footprint of the cross-multiplication component of the AMI digital correlator has been reduced to a very small fraction of the overall design (Table 2).

Vector Accumulator Per-frequency-channel visibility matrices, which have been integrated over 1024 samples, must be further time-averaged to reduce output data-rate to an acceptable level. This long-term (\sim second) averaging uses external Quad Data Rate (QDR) memory for data storage, at the end of which data are transmitted via a 1 Gb Ethernet interface to a data acquisition and storage server.

3.6 FPGA Floorplanning

A significant hurdle in the deployment of the new AMI correlator was achieving timing-closure of the firmware designs at the required 312.5 MHz FPGA clock rate. Ultimately, this was achieved by constraining the placement of most of the \sim 1000 multiplier cores and \sim 400 RAM blocks in the design to small regions of the FPGA chip. In the case of the FFT module, a pair of which comprise the majority of the signal processing in the firmware, a set of Python scripts was used to compute the number of multiplier and memory cores needed by the module, and to automatically place the cores at specific sites designed to maximise timing performance. To the knowledge of the authors, the AMI design represents the fastest clocked ROACH2 in any deployed packetised-correlator system, with the largest instantaneous processed bandwidth per board. We note, however, that the AMI firmware resource requirements (Table 2) are modest when compared to some other deployments, such as Primiani et al. (2016). Experience with newer FPGA platforms also suggests that many of the timing issues encountered during development of the AMI correlator may be unique to the ROACH2’s Virtex 6 FPGA.

4 CONTROL SYSTEM

The control systems for the two arrays are independent, and are described in detail in Zwart et al. (2008). In each case, the only modifications needed for the new correlator have been replacement of the correlator microprocessor by the data acquisition and control server, situated in the digital rack and visible in Figure 7, and replacement of the Sun-Microsystems workstation used for telescope and observation control by a Linux desktop system. Apart from the data acquisition components described below, the rest of the control software, written in C for the Solaris operating system, has simply been recompiled for the Linux platform.

4.1 Data Acquisition

The Linux server used for correlator control and data acquisition is a standard single-socket machine, configured with 20 TB of 10-

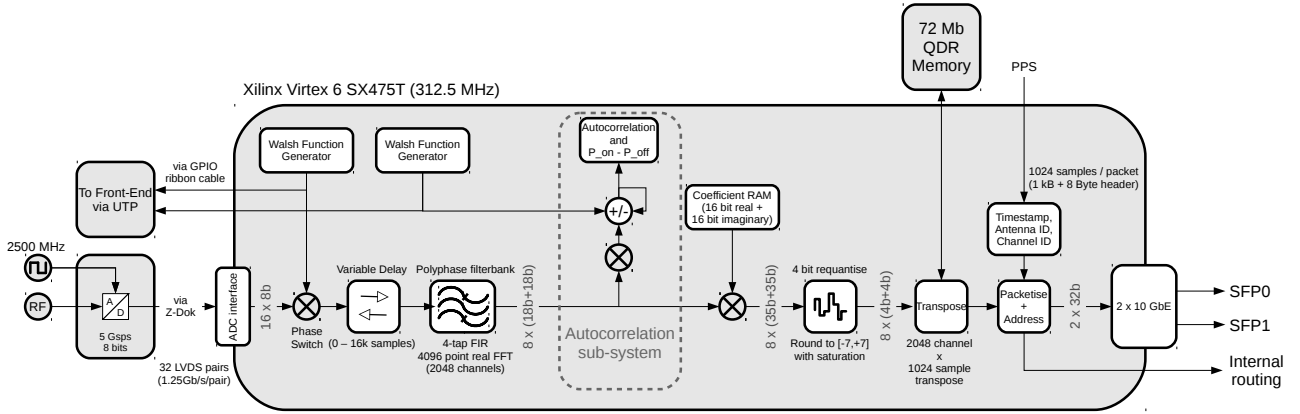


Figure 9. A schematic depiction of the AMI F-Engine firmware. Two such F-Engines are instantiated on each ROACH2 FPGA.

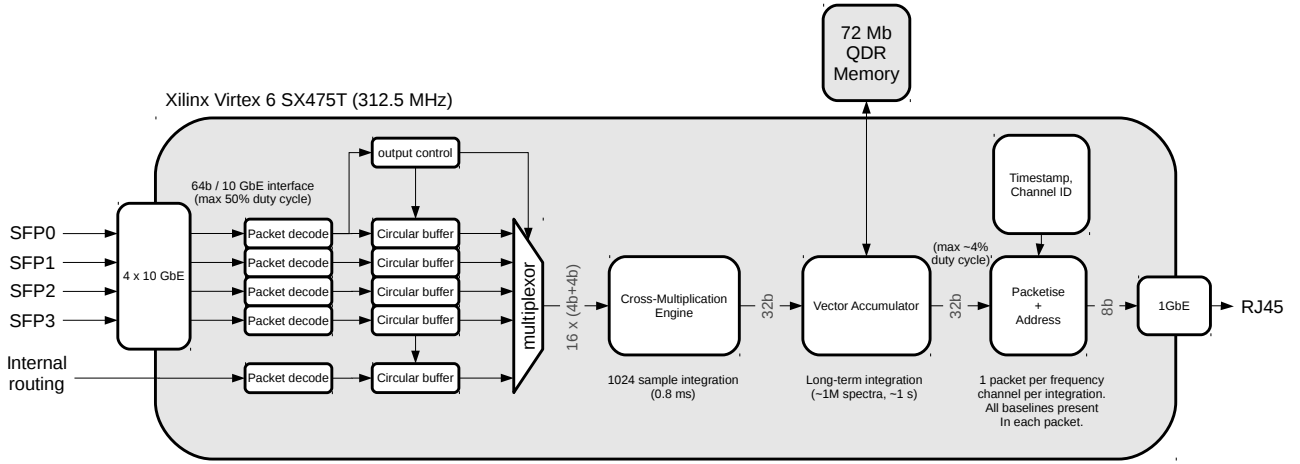


Figure 10. A schematic depiction of the AMI X-Engine firmware.

Table 2. FPGA resource utilisation in the AMI SA correlator design. The AMI LA design uses slightly fewer resources, owing to the smaller number of antennas requiring correlation in the X-Engine stage of the design.

	DSP slices	BRAMs (36 kB)	LUTs	Flip Flops
Coarse Delay (dual-band)	0	16	2024	2314
FIR filter (dual-band)	128	66	3952	11890
FFT (dual-band)	736	112	38206	100748
X-Engine input buffers	0	81	3162	10747
X-Engine	98	18	2002	4722
Total Design	1203	409	105068	166051
FPGA Capacity	2016	1064	297600	595200
% Utilisation	60	38	35	28

cal disc storage to accommodate several months of raw data. A 10 Gbps network interface has been added for communication with the ROACH2 boards, with the on-board 1 Gbps interface used for exchanging data with the telescope control system over a local Ethernet connection. This machine also acts as the DHCP/NFS server for the ROACH2 private network.

A suite of Python scripts to configure the ROACH2 boards and perform data collection has been developed for AMI, and is available on github⁷. There are separate scripts for programming and initialising the ROACH2 boards, and for computing the equalisation

coefficients for requantization (Section 3.3). Acquisition scripts deal with applying coarse path compensation delays to the digital data streams, exchanging metadata and visibilities with the telescope control machine and recording visibility data from the correlator at a programmable dump rate (currently every 0.86 s) in HDF5⁸ format on the server's local discs. The arrangement for monitoring the relative system temperature of each antenna using amplitude-modulated injected noise referred to earlier and described in (Zwart et al. 2008) is retained, and samples of the demodulated noise,

⁷ https://github.com/jack-h/ami_correlator_sw.git

⁸ <https://www.hdfgroup.org/HDF5/>

Table 3. Typical AMI data rates.

	Small Array	Large Array
Sampling rate (Mbps)	16.8	11.0
Recording rate (GBph)	8	5.2
8-hour observation (GB)	64	42

derived from the autocorrelation measurements, are stored with the data for use during calibration processing.

Important features of the software design include the use of a standard configuration file, which allows the same code to be used for both arrays, and a REmote Dictionary Service (Redis⁹) key-value data store to share all data among the independently executed processes. A comprehensive set of utilities for monitoring the state of the correlator and visualising the datasets is also included in the AMI software package.

Independently of data acquisition from the correlator, but at a similar sampling rate, metadata such as antenna pointing and instrumental monitoring information (e.g. cryostat temperatures, AGC readings) are collected by the telescope control machine, and exchanged with the correlator server over the local Ethernet.

4.2 Data Handling

The HDF5 format data files, one per observation, include correlated visibility data, noise-injection measurements and metadata samples, and accumulate on discs local to the correlator control server. Typical raw data rates are shown in Table 3. These files are the primary data products from the telescope and contain uncalibrated, full spectral resolution data.

After each observation the HDF5 dataset is converted to FITS-IDI¹⁰ format for calibration and further processing using standard reduction packages, and is transferred from the observatory to a data repository at the Cavendish Laboratory. The conversion script is based on the `pyFitsidi`¹¹ Python module, and also performs the following operations:

- removal of artefacts in certain frequency channels.
- removal of the channel readout delay phase shift.
- fringe rotation to remove the astronomical path difference over each sub-band.
 - amplitude correction for sampling rate (significant for the LA).
 - amplitude calibration for system temperature using the noise-injection system.
 - flagging for instrumental problems (e.g. AGCs, cryostats, pointing, shadowing).
 - optional binning of frequency channels, to reduce the dataset size.

In parallel to the above, the raw data samples received from the correlator are binned to 8×0.625 GHz frequency channels by a new task added to the real-time software on the telescope control system, and are recorded in legacy format for quick-look using the in-house software tool, REDUCE. These 8 channels are chosen to correspond to those used by the old analogue correlator, and availability of these datasets has proved invaluable for data evaluation and diagnostics throughout the development process and during routine operation.

5 DATA REDUCTION

The FITS-IDI files contain visibility data to which instrumental flagging and some initial calibration steps have been applied, and these can now be imported to the standard reduction packages for radio interferometric data such as AIPS¹² and CASA¹³ for full calibration, radio-frequency interference (RFI) flagging and map-making. A generalised pipeline script has been developed to process these datasets within the CASA framework, and deals with:

- RFI flagging, both broad-band and at full frequency channel resolution.
- bandpass (amplitude and phase) and flux calibration, using observations of standard calibrators.
- phase calibration, using data from interleaved visits to nearby high flux density, unresolved radio sources.
- imaging.

Primary beam correction and mosaicing is currently done in AIPS due to difficulties in importing new primary beam models into CASA.

This is the recommended route for processing AMI data for science. The 8-channel legacy format datasets described above are also available and can be processed using the in-house REDUCE software to provide useful results for some projects, but lose the benefits of the full spectral resolution that the new correlator provides.

6 COMMISSIONING RESULTS

6.1 RFI rejection

One of the main motivators for constructing a correlator with such improved spectral resolution was to be able to identify and remove radio-frequency interference (RFI), which is typically confined to narrow frequency bands. RFI is generally worse at low elevation, due to the location of geostationary satellites. With the old correlator, this limited the useable declination range, particularly for the SA with its larger primary beam. RFI flagging levels on the SA reached $\approx 85\%$ at $\delta \approx 2^\circ$; in practice observations were rarely made below $\delta = 20^\circ$. In Figure 11 we show part of an observation of a low-declination ($\approx 0^\circ$) field on both the LA and the SA with the new correlator; the RFI is clearly localized in time- and frequency-space and can be easily removed. At $\delta \approx 0^\circ$ the RFI flagging level is now $\approx 30\%$ on the SA and $\approx 20\%$ on the LA (due to the smaller primary beam).

6.2 Dynamic range

The uneven lag spacings of the old correlator limited the dynamic range, since maps of fields near high flux density sources would be contaminated by artefacts introduced by the correlator. The digital correlator has removed this issue and consequently the dynamic range of the telescope, defined as the ratio of the brightest believable flux to brightest non-believable flux on the map, has increased from ~ 100 to ~ 1000 . Figure 12 and Figure 13 show example maps of a bright source produced on the LA and the SA, in comparison to similar maps using old correlator data. The improved dynamic

⁹ <http://redis.io/>

¹⁰ <http://www.nrao.edu/aips/FITSIDI.pdf>

¹¹ <http://telegraphic.github.io/pyfitsidi>

¹² <http://www.aips.nrao.edu/>

¹³ <http://casa.nrao.edu>

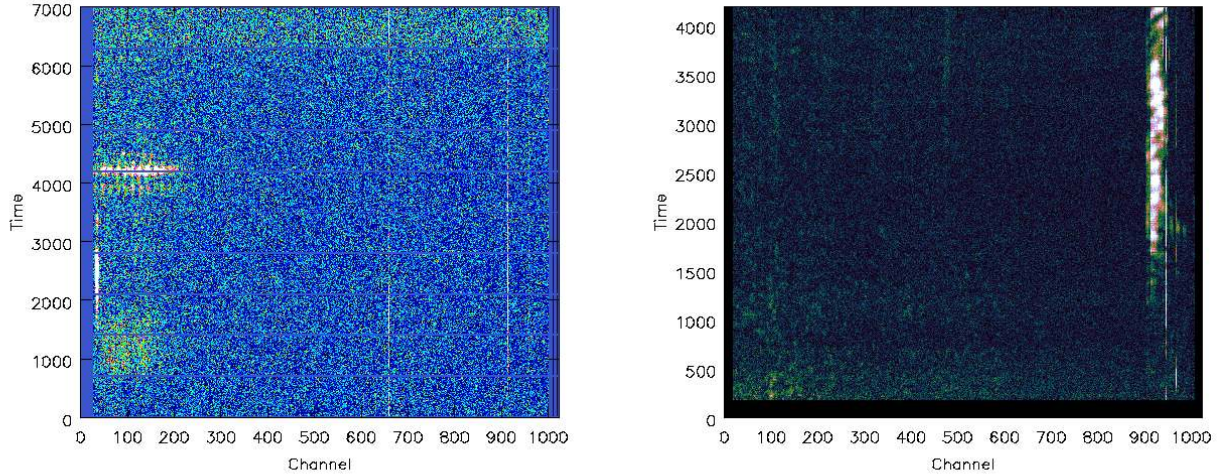


Figure 11. ‘Waterfall’ plots of a low-declination ($\approx 0^\circ$) field observed with the LA (SA) on the left (right), showing time against channel for the lower (upper) half of the bandwidth on a single baseline. The RFI can be clearly seen as bright signals localised in frequency and time and is easily removed using automated algorithms.

range of the new AMI system is enabling, for example, the observation of galaxy clusters containing bright radio galaxies, which were previously excluded from cluster samples potentially introducing unquantifiable biases.

6.3 Scientific Collaboration

AMI has been operating routinely with the upgraded correlator since January 2016. The current scientific programmes include:

- SZ observations of galaxy clusters, for example extending the previous follow-up of *Planck* galaxy clusters (Perrott et al. 2015) to include clusters previously excluded for difficult radio source environments
- Source surveys, extending the 10C source counts (e.g. Whitam et al. 2016) to lower flux densities and adding to the multi-wavelength legacy datasets available in fields such as Stripe 82 (e.g. Ivezić et al. 2007 and COSMOS (e.g. Scoville et al. 2007))
- Monitoring of variable and transient radio sources such as supernovae, GRBs, quasars and X-ray binaries
- Observations of supernova remnants with reported anomalous microwave emission (AME) detections at low resolution (Génova-Santos et al. 2017, Onić et al. 2017), to test for the presence of AME at the angular scales measured by AMI.

The first results using data taken with the upgraded correlator have been published (Muñoz-Darias et al. 2017, Mooley et al. 2017); more are currently in preparation.

The telescope operates under the terms of a Scientific Collaboration Agreement involving the University of Cambridge and the University of Manchester, but welcomes external collaborators, and expressions of interest in observing with AMI may be addressed to ami-request@mrao.cam.ac.uk. AMI is unusual in offering an automatically triggered rapid-response mode which has been exploited in studies of source variability and transients (Staley et al. (2013)). For reference, a summary of the technical data for AMI observations is given in Table 4.

Table 4. Summary of AMI technical data.

	Small Array	Large Array
Antenna diameter (m)	3.7	12.8
Antenna efficiency	0.75	0.67
Number of antennas	10	8
Number of baselines	45	28
Baseline lengths (m)	5–20	18–110
Primary beam (@15.5 GHz)	20′	5′
Synthesized beam	$\approx 3'$	$\approx 30''$
Flux sensitivity ($\text{mJy s}^{1/2}$)	30	3
Declination range ($^\circ$)	> -15	> -20
Elevation limit ($^\circ$)	+20	+5
Observing frequency (GHz)	13.1–17.9	
Number of channels	4096	
Channel bandwidth (MHz)	1.22	
System temperature (K)	25	
Polarisation measurement	1+Q	

7 CONCLUSIONS

In this paper we have presented a new digital correlator system for the AMI telescope, which replaces the telescope’s previous analogue lag-correlator. This real-time FX correlator is implemented using the popular ‘‘packetised correlator’’ architecture, in which interconnect between processors is provided by commercial Ethernet switches. Processing in the new correlator is performed on CASPER open-source hardware; FPGA-based ROACH2 platforms are used to digitise, filter into 1.2 MHz-wide channels, and correlate a pair of down-converted sub-band signals at 5 Gsps, providing a usable RF band of 13.1–17.9 GHz. This wideband performance has been achieved by clocking the signal processing pipelines on the ROACH2’s Xilinx Virtex 6 FPGAs at 312.5 MHz, requiring optimisation and floorplanning of the firmware design.

The improvement in performance of the new instrument when compared to the original analogue XF correlator is evident; the new correlator achieves over an order of magnitude improvement in imaging dynamic range, and is far more effective at operating in the presence of RFI – particularly interference at low-declination ob-

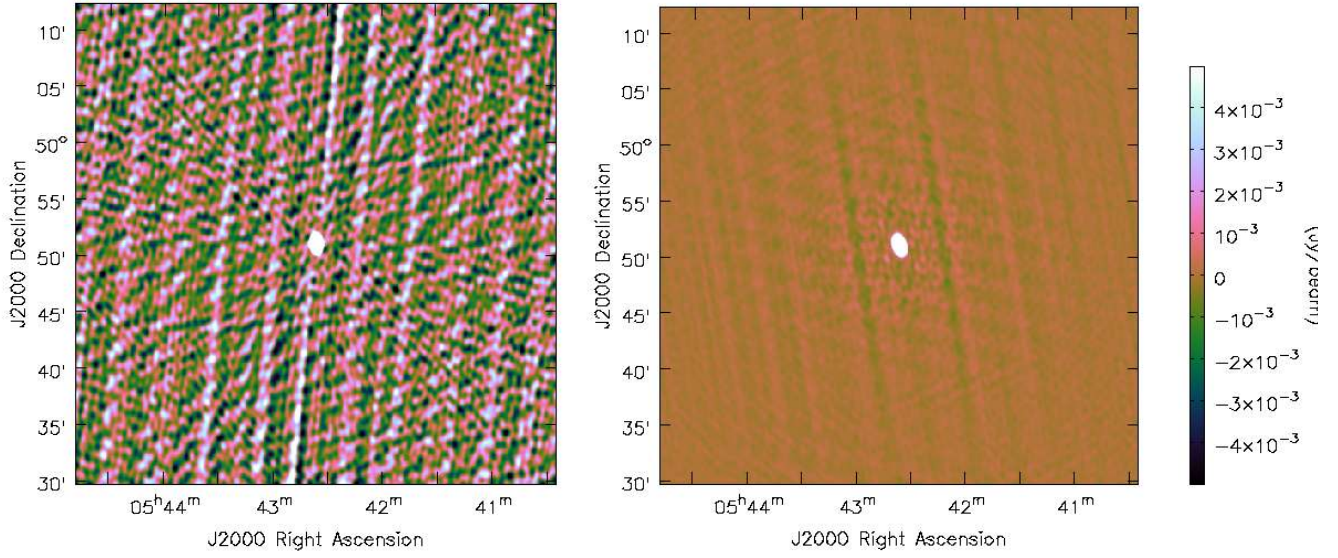


Figure 12. An LA observation of the bright, unresolved source 3C147 using the old (new) correlator on the left (right). The colour scales are the same, and have been truncated to show the systematic residuals around the source. These are at most 9.3 (1.1) mJy beam^{-1} , giving a dynamic range of 260 (2300).

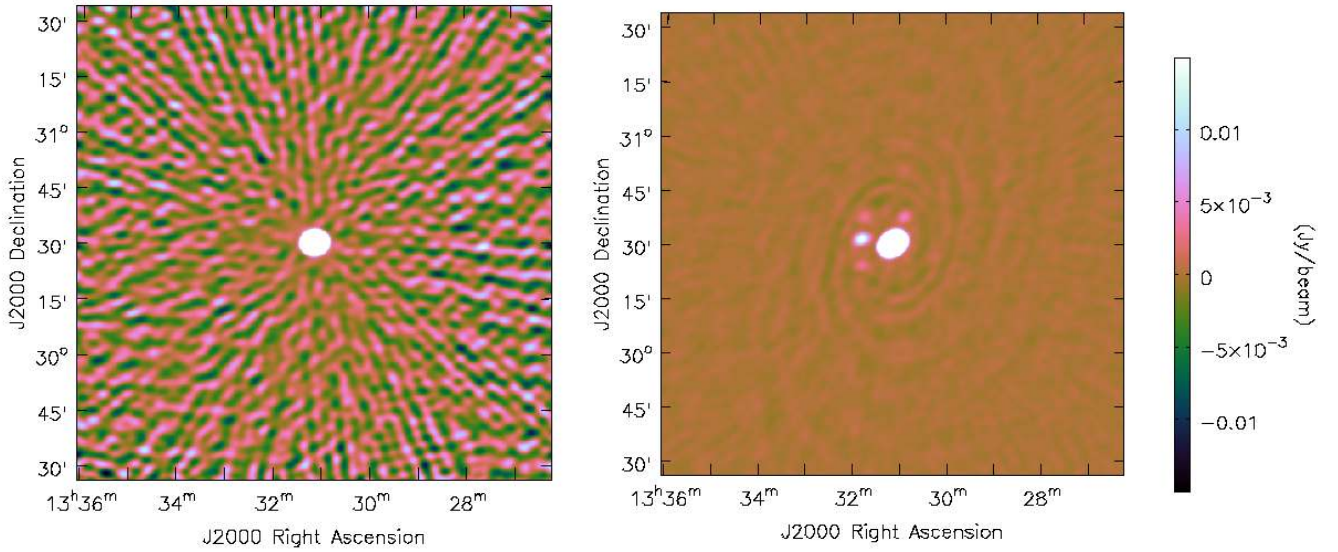


Figure 13. An SA observation of the bright, unresolved source 3C286 using the old (new) correlator on the left (right). The colour scales are the same, and have been truncated to show the systematic residuals around the source. These are at most 15 (1.2) mJy beam^{-1} , giving a dynamic range of 230 (3000). The integration time for the new correlator map was much longer, in order for the thermal noise to fall below the level of the systematic errors. The triple of sources visible to the east of 3C286 in the new correlator map is real (the central source is FIRST J133148.5+303147); the source to the north of 3C286 is also real (FIRST J133052.9+303807). They are all under the level of the residuals in the old correlator map.

servations from geostationary satellites – owing to superior spectral resolution. As a result of the improved performance, a new scientific collaboration has been established for the operation of AMI, and a broad range of observations are currently being performed with the telescope, which is now highly subscribed.

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REFERENCES

- Bellanger M., Bonnerot G., Coudreuse M., 1976, *IEEE Transactions on Acoustics, Speech, and Signal Processing*, 24, 109
- Foley A. R., et al., 2016, *Monthly Notices of the Royal Astronomical Society*, 460, 1664
- Génova-Santos R., et al., 2017, *MNRAS*, 464, 4107
- Green D. A., 2011, *Bulletin of the Astronomical Society of India*, 39, 289

- Harris C., Haines K., 2011, Publications of the Astronomical Society of Australia, 28, 317
- Hickish J., 2014, PhD thesis, University of Oxford
- Hickish J., et al., 2016, Journal of Astronomical Instrumentation
- Ivezić Ž., et al., 2007, *AJ*, 134, 973
- Klages P., Bandura K., Denman N., Recnik A., Sievers J., Vanderlinde K., 2015, preprint, ([arXiv:1503.06203](https://arxiv.org/abs/1503.06203))
- Kocz J., et al., 2015, *Journal of Astronomical Instrumentation*, 04, 1550003
- Mooley K. P., et al., 2017, *MNRAS*, 467, L31
- Muñoz-Darias T., et al., 2017, *MNRAS*, 465, L124
- Nita G., Hickish J., MacMahon D., Gary D., 2016, JAI
- Onić D., Urošević D., Leahy D., 2017, *AJ*, 153, 32
- Parsons A., et al., 2008, Publications of the Astronomical Society of the Pacific, 120, pp. 1207
- Patel N. A., Wilson R. W., Primiani R., Weintroub J., Test J., Young K., 2014, *Journal of Astronomical Instrumentation*, 03, 1450001
- Perrott Y. C., et al., 2015, *A&A*, 580, A95
- Price D. C., 2016, arXiv preprint arXiv:1607.03579
- Primiani R., et al., 2016, JAI
- Scoville N., et al., 2007, *ApJS*, 172, 1
- Staley T. D., et al., 2013, *MNRAS*, 428, 3114
- Sunyaev R. A., Zeldovich Y. B., 1970, *Ap&SS*, 7, 3
- Whittam I. H., Riley J. M., Green D. A., Davies M. L., Franzen T. M. O., Rumsey C., Schammel M. P., Waldram E. M., 2016, *MNRAS*, 457, 1496
- Zwart J. T. L., et al., 2008, *MNRAS*, 391, 1545
- de Souza L., Bunton J., Campbell-Wilson D., Cappallo R., Kincaid B., 2007, in Field Programmable Logic and Applications, 2007. FPL 2007. International Conference on. pp 62–67, doi:10.1109/FPL.2007.4380626

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