

 Open access • Journal Article • DOI:10.3847/2041-8213/AB0379

A mass-dependent slope of the galaxy size-mass relation out to $z \sim 3$: further evidence for a direct relation between median galaxy size and median halo mass

— [Source link](#) 

Lamiya Mowla, Arjen van der Wel, Pieter G. van Dokkum, Tim B. Miller

Institutions: Max Planck Society

Published on: 15 Jan 2019 - arXiv: Astrophysics of Galaxies

Topics: Stellar mass, Galaxy and Mass ratio

Related papers:

- [Hyper Suprime-Cam Subaru Strategic Program: A Mass-dependent Slope of the Galaxy Size–Mass Relation at \$z < 1\$](#)
- [The growth of brightest cluster galaxies and intracluster light over the past 10 billion years](#)
- [The Rest-frame Optical \(900 nm\) Galaxy Luminosity Function at \$z \sim 4\$ -7: Abundance Matching Points to Limited Evolution in the \$M_{\text{STAR}}/M_{\text{HALO}}\$ Ratio at \$z \geq 4\$](#)
- [The Evolution of the Galaxy Stellar Mass Function at \$z = 4\$ -8: A Steepening Low-mass-end Slope with Increasing Redshift](#)
- [Half-mass radii for \$\sim 7,000\$ galaxies at \$1.0 < z < 2.5\$: most of the evolution in the mass-size relation is due to color gradients](#)

Share this paper:    

View more about this paper here: <https://typeset.io/papers/a-mass-dependent-slope-of-the-galaxy-size-mass-relation-out-14sq4as466>



A Mass-dependent Slope of the Galaxy Size–Mass Relation out to $z \sim 3$: Further Evidence for a Direct Relation between Median Galaxy Size and Median Halo Mass

Lamiya Mowla¹ , Arjen van der Wel^{2,3} , Pieter van Dokkum¹ , and Tim B. Miller¹

¹Astronomy Department, Yale University, New Haven, CT 06511, USA; lamiya.mowla@yale.edu

²Sterrenkundig Observatorium, Universiteit Gent, Krijgslaan 281 S9, B-9000 Gent, Belgium

³Max-Planck-Institut für Astronomie, Königstuhl 17, D-69117, Heidelberg, Germany

Received 2018 December 21; revised 2019 January 26; accepted 2019 January 30; published 2019 February 13

Abstract

We reassess the galaxy size–mass relation out to $z \sim 3$ using a new definition of size and a sample of $>29,000$ galaxies from the 3D-*HST*, CANDELS, and COSMOS-DASH surveys. Instead of the half-light radius r_{50} we use r_{80} , the radius containing 80% of the stellar light. We find that the r_{80} – M_* relation has the form of a broken power law, with a clear change of slope at a pivot mass M_p . Below the pivot mass the relation is shallow ($r_{80} \propto M_*^{0.15}$); above it, it is steep ($r_{80} \propto M_*^{0.6}$). The pivot mass increases with redshift, from $\log(M_p/M_\odot) \approx 10.2$ at $z = 0.4$ to $\log(M_p/M_\odot) \approx 10.9$ at $z = 1.7$ – 3 . We compare these r_{80} – M_* relations to the M_{halo} – M_* relations derived from galaxy–galaxy lensing, clustering analyses, and abundance matching techniques. Remarkably, the pivot stellar masses of both relations are consistent with each other at all redshifts, and the slopes are very similar both above and below the pivot when assuming $M_{\text{halo}} \propto r_{80}^3$. The implied scaling factor to relate galaxy size to halo size is $r_{80}/R_{\text{vir}} = 0.047$, independent of stellar mass and redshift. From redshift 0 to 1.5, the pivot mass also coincides with the mass where the fraction of star-forming galaxies is 50%, suggesting that the pivot mass reflects a transition from dissipational to dissipationless galaxy growth. Finally, our results imply that the scatter in the stellar-to-halo mass is relatively small for massive halos (~ 0.2 dex for $M_{\text{halo}} > 10^{12.5} M_\odot$).

Key words: galaxies: evolution – galaxies: halos – galaxies: high-redshift – galaxies: structure

1. Introduction

The size distribution of galaxies holds clues to their assembly history and the relationship with their dark matter halos (Mo et al. 1998; Kravtsov 2013; Jiang et al. 2018). The sizes of galaxies are known to vary with stellar mass, star formation rate, and redshift, and have been studied extensively (e.g., Kormendy 1977; Shen et al. 2003; Ferguson et al. 2004; Trujillo et al. 2006; Elmegreen et al. 2007; Williams et al. 2010; Mosleh et al. 2012; Carollo et al. 2013; Ono et al. 2013; Bernardi et al. 2014; van der Wel et al. 2014; Navarro et al. 2017; Kravtsov et al. 2018; Mowla et al. 2018, among many others).

One key result of these studies is that the two main classes of galaxies, star-forming and quiescent, follow very different size–mass relations. Hence it has been common practice to describe the size–mass distribution of galaxies separately for the two classes. It is usually defined by a single power-law relation for each sub-population, with quiescent galaxies having a steeper relation than star-forming ones (e.g., Shen et al. 2003; van der Wel et al. 2014). The interpretation of these results is a topic of debate; one possibility is that star-forming galaxies build up their stellar populations at all radii, whereas quiescent galaxies mostly grow inside-out through accretion (e.g., van Dokkum et al. 2015).

In this Letter we revisit the form of the size–mass relation out to $z = 3$, using a large sample and an alternative size definition. This study is motivated by the availability of a new, large sample of distant galaxies with *Hubble Space Telescope* (*HST*)-measured sizes out to $z = 3$ (Mowla et al. 2018), which extends to higher masses than previous studies (van der Wel et al. 2014).

We find that the size–mass distribution of all galaxies is not well fit by a single power law, but requires a change in slope at a pivot

mass. We compare this to stellar-to-halo mass (SMHM) relations from the literature, assuming a conversion from size to virial radius (e.g., Leauthaud et al. 2012; Moster et al. 2013; Behroozi et al. 2018). Our study extends earlier work at low redshift that found a steepening of the size–mass distribution at the high-mass end for disk-dominated galaxies (e.g., Shen et al. 2003; Dutton et al. 2011), and theoretical work that suggested a constant scaling between the half-light radius and the virial radius of galaxies (Kravtsov 2013; Huang et al. 2017; Genel et al. 2018; Jiang et al. 2018; Somerville et al. 2018). We assume a flat Λ CDM cosmology with parameters $\Omega_m = 0.308$, $\Omega_b = 0.049$, $h = H_0/(100 \text{ km s}^{-1} \text{ Mpc}^{-1}) = 0.677$, $\sigma_8 = 0.823$, and $n_s = 0.96$ compatible with *Planck* constraints (Planck Collaboration et al. 2016).

2. Data

2.1. Galaxy Sample

The data set that we use is described in Mowla et al. (2018) and consists of the combination of two distinct samples. The first is from the CANDELS/3D-*HST* surveys. Sizes of over 28,000 galaxies with $M_* > 10^9 M_\odot$ at $0 < z < 3$ were measured by van der Wel et al. (2014) from the 0.22 deg^2 CANDELS (Koekemoer et al. 2011) imaging in the H_{160} , J_{125} , and I_{814} bands. Spectroscopic and photometric redshifts, stellar masses, and rest-frame properties were measured by Skelton et al. (2014) using the extensive 3D-*HST* multi-wavelength data.

The area of the CANDELS/3D-*HST* fields is insufficient to properly sample the massive end of the luminosity function. This situation has been mitigated by the completion of the COSMOS-DASH survey, which tripled the area surveyed by *HST* in the near-infrared (near-IR). COSMOS-DASH has enabled us to extend the size–mass study to higher masses at $1.5 < z < 3.0$ (Mowla et al. 2018). Sizes of 162 galaxies with

$M_* > 2 \times 10^{11} M_\odot$ at $z > 1.5$ were measured from H_{160} COSMOS-DASH imaging (0.66 deg^2) and of 748 galaxies at $z < 1.5$ from I_{814} ACS-COSMOS imaging (1.7 deg^2) (Koekemoer et al. 2007; Massey et al. 2010). Photometric redshifts, stellar masses, and rest-frame colors were taken from the UltraVISTA catalog (Muzzin et al. 2013).

In both surveys the sizes of galaxies were measured by single-component Sèrsic profile fits to 2D light distributions using GALFIT (Peng et al. 2010), with a correction for redshift-dependent color gradients. Details are described in van der Wel et al. (2014) and Mowla et al. (2018). The two data sets have been combined carefully, verifying that there are no detected systematic differences between them. The combined size–mass distribution is the largest data set with the largest ranges in stellar mass and redshift currently available, and is detailed in Mowla et al. (2018). In this Letter we only analyze data where we are mass-complete. The lower bounds of the stellar mass limits correspond to the mass-completeness limits down to which van der Wel et al. (2014) determined structural parameters for star-forming and quiescent galaxies with good fidelity.

2.2. Galaxy Size Definition

High-redshift galaxies are typically modeled by single-component Sèrsic profiles that describe the structure of a galaxy with the half-light radius r_{50} , the radius containing 50% of light, and the Sèrsic index n , a measure of form of the light profile. On average, at a given stellar mass quiescent galaxies have a higher Sèrsic index and a smaller half-light radius than star-forming galaxies. As demonstrated in a companion Letter (Miller et al. 2019), these two effects conspire such that the size difference between star-forming and quiescent galaxies nearly disappears when using r_{80} , the radius containing 80% of the stellar light. The size–mass relation is also tighter for this definition of radius by approximately 0.06 dex (see Miller et al. 2019). Physically, this size definition is a better measure of the total baryonic extent, and it is in a regime where dark matter begins to dominate the mass. For a typical galaxy with $M_* \sim 5 \times 10^{10} M_\odot$ and dark matter halo mass $\sim 10^{12} M_\odot$, the median fraction of dark matter contained within r_{50} is 35%–55%, while that within r_{80} is between 60% and 80%.

Following Miller et al. (2019), we calculate r_{80} using the following relation:

$$\frac{r_{80}}{r_{50}} = 0.0012n^3 - 0.0123n^2 + 0.5092n + 1.2646, \quad (1)$$

with r_{50} the half-light radius and n the Sèrsic index.

3. The Size–Mass Relation

3.1. Broken Power-law Fit

The r_{80} –mass distributions of all galaxies with $\log(M_*/M_\odot) > 9$ in six bins of redshift are shown in Figure 1. The visible gaps in the distributions mark the points where the CANDELS sample (van der Wel et al. 2014) transitions to the high-mass ULTRAVISTA/COSMOS sample (Mowla et al. 2018). The median sizes of galaxies in mass bins are overplotted on the size–mass distribution, which are color-coded by the fraction of galaxies which are quiescent; rest-frame $U - V$ and $V - J$ color space was used to separate galaxies into star-forming and quiescent. The error bars on median sizes are calculated as biweight scale divided by $\sqrt{N - 1}$, where N is the number of

galaxies in each stellar mass bin. Visual inspection of the median size–mass relation shows that at the low-mass end the relation has a shallow slope, while at the high-mass end the relation steepens after a characteristic pivot point. Hence, we fit a smoothly broken power law to the median size–mass relation of the form

$$r_{80}(M_*) = r_p \left(\frac{M_*}{M_p} \right)^\alpha \left[\frac{1}{2} \left(1 + \left(\frac{M_*}{M_p} \right)^\delta \right) \right]^{(\beta - \alpha)/\delta}, \quad (2)$$

where M_p is the pivot stellar mass at which the slopes change, r_p is the radius at the pivot stellar mass, α is the slope at the low-mass end, β is the slope at the high-mass end, and δ is the smoothing factor. We set the smoothing factor to $\delta = 6$ to reduce degeneracy between δ and the slopes. We fit Equation (2) to the median sizes at each redshift bin, using the “trust region reflective” algorithm as implemented in `curvefit` of `scipy`. The parameters of the best-fitting relations for all redshift bins are given in Table 1.

The fits are generally excellent, with reduced χ^2 values ranging between 0.95 and 1.8. The most stable results are obtained in the redshift range $0.5 < z < 2.0$, as there are >5000 galaxies in each redshift bin with a large dynamic range in mass. At the lowest redshifts ($z < 0.5$) the COSMOS field does not have sufficient volume to properly sample the full distribution; this may affect the characteristic pivot mass measurement. A similar problem arises at $z > 2$, where the pivot stellar mass is high and we have a relatively low number of galaxies in the relevant mass range.

3.2. Redshift Evolution of the Size–Stellar Mass Relation

The parameters of the best-fit broken power-law function are shown as a function of redshift in the left panel of Figure 2. The best-fitting functions are shown in the right panel, and are also overplotted in Figure 1. We find that the pivot stellar mass decreased with cosmic time, going from $M_p = 8 \times 10^{10} M_\odot$ at $z \sim 1.72$ to $M_p = 1.4 \times 10^{10} M_\odot$ at $z \sim 0.25$. The evolution in pivot stellar mass appears to flatten off between $z \sim 1.5$ and $z \sim 3$; however, further study is required to investigate whether or not this is a physical phenomenon or due to small number statistics of high-mass galaxies at $z > 2$. In contrast to the pivot mass itself, the radius at the pivot mass increased with time, from $r_p = 5.3 \text{ kpc}$ at $z \sim 2.75$ to $r_p = 8.6 \text{ kpc}$ at $z \sim 0.25$.

The slope of the size–mass relation at $M_* < M_p$ is approximately constant at $\alpha \approx 0.16$ while it is $\beta \approx 0.60$ at $M_* > M_p$. We note here that the slope of the low-mass end is similar to the slope of a single power-law fit to the sample of star-forming galaxies, while that of the high-mass galaxies is similar to a single power-law slope of quiescent galaxies (see Mowla et al. 2018).

4. Halo-to-stellar-mass Relation

4.1. Calculating Halo Mass

The functional form of the size–mass relation is reminiscent of the form of the SMHM relation: this relation also has different slopes in different mass regimes with an inflection point. In the SMHM relation, the inflection point is where galaxy formation is maximally efficient in the sense that the largest fraction of baryons is in stars (Behroozi et al. 2010). This superficial similarity motivates us to examine the

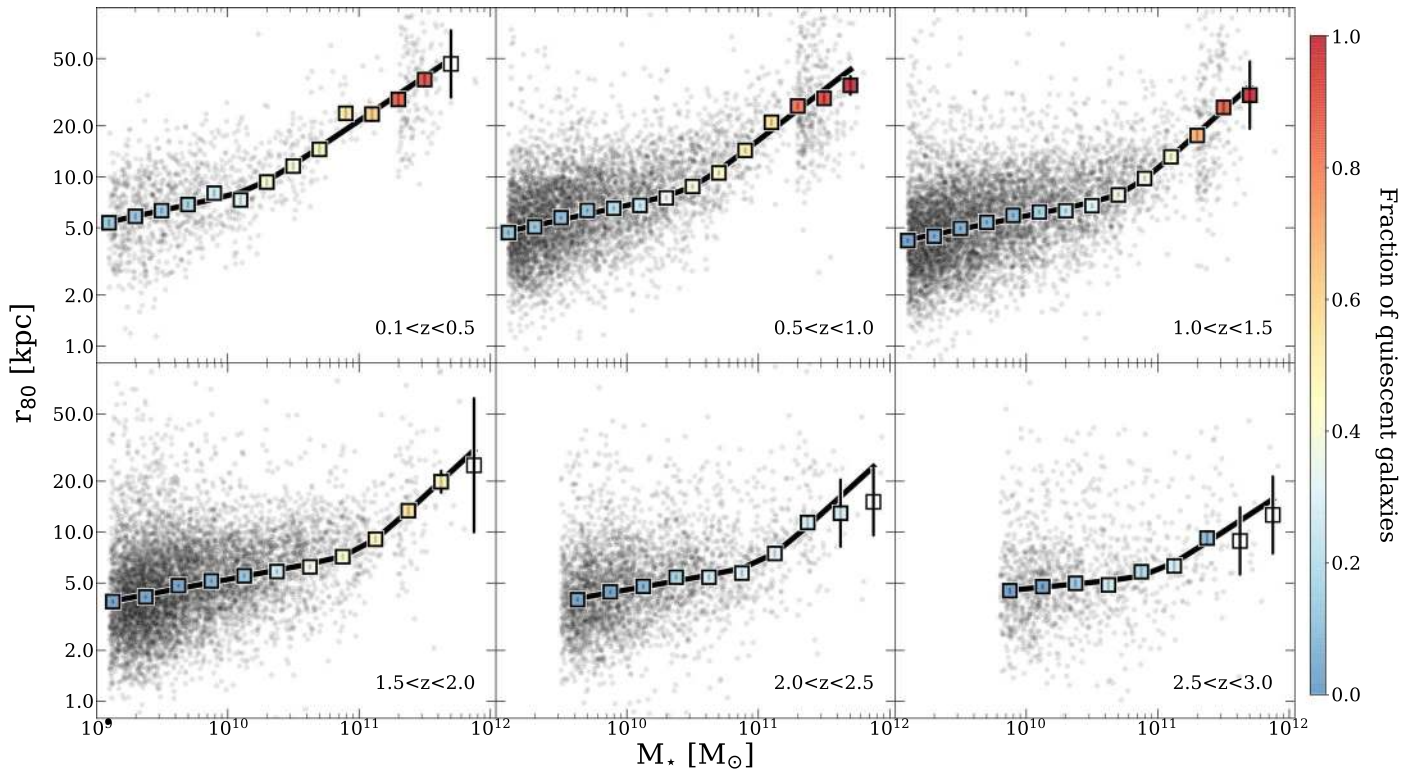


Figure 1. Size–stellar mass distribution of galaxies at $0 < z < 3$ from van der Wel et al. (2014) and Mowla et al. (2018). The squares show the median of r_{80} in bins of $\log(M_*/M_\odot)$, color-coded by the fraction of quiescent galaxies in the bin; rest-frame $U - V$ and $V - J$ color space was used to separate galaxies into star-forming and quiescent. Unfilled squares represent bins with less than 15 galaxies. Smoothly broken power-law fits given by Equation (2) to the median size–mass relation are shown by the black lines.

Table 1
Best-fit Parameters of Smoothly Broken Power-law Fit to the Size–Mass Relation, Given in Equation (2)

z	r_{80}				r_{50}			
	r_p (kpc)	$\log(M_p/M_\odot)$	α	β	r_p (kpc)	$\log(M_p/M_\odot)$	α	β
0.37	8.6 ± 0.7	10.2 ± 0.1	0.17 ± 0.03	0.50 ± 0.03	3.8 ± 0.3	10.3 ± 0.1	0.09 ± 0.03	0.37 ± 0.03
0.79	8.7 ± 0.5	10.5 ± 0.1	0.17 ± 0.02	0.61 ± 0.04	4.0 ± 0.4	10.7 ± 0.2	0.10 ± 0.02	0.45 ± 0.09
1.24	8.3 ± 0.3	10.8 ± 0.1	0.16 ± 0.01	0.69 ± 0.06	4.2 ± 0.4	11.1 ± 0.2	0.13 ± 0.01	0.53 ± 0.17
1.72	7.6 ± 0.7	10.9 ± 0.1	0.15 ± 0.01	0.62 ± 0.19	3.7 ± 0.8	11.1 ± 0.5	0.11 ± 0.03	0.50 ± 0.29
2.24	6.5 ± 0.7	11.0 ± 0.2	0.14 ± 0.02	0.53 ± 0.17	3.1 ± 0.5	11.0 ± 0.3	0.11 ± 0.02	0.42 ± 0.25
2.69	5.3 ± 0.4	10.8 ± 0.2	0.05 ± 0.03	0.34 ± 0.09	2.8 ± 0.4	10.9 ± 0.3	0.06 ± 0.04	0.38 ± 0.20

hypothesis that the upturn in the stellar size–stellar mass relation above the pivot stellar mass is simply a reflection of the downturn in the SMHM relation above its pivot halo mass.

We test this by adopting a constant ratio between galaxy size and the virial radius of the halo: $R_{\text{vir}} = r_{80}/\gamma$. We define the halo virial mass and virial radius within a spherical overdensity Δ_{vir} times the critical density ρ_{crit} :

$$M_{\text{halo}} = \frac{4\pi}{3} \Delta_{\text{vir}} \rho_{\text{crit}} R_{\text{vir}}^3, \quad (3)$$

where Δ_{vir} is from Bryan & Norman (1998).

This allows us to express median galaxy radii, r_{80} , in terms of median halo masses, by choosing an appropriate value for the proportionality constant γ . We fit for γ by minimizing the difference between the SMHM relation that we derive in the lowest redshift bin and the relation from Leauthaud et al. (2012) at $0.2 < z < 0.48$. Leauthaud et al. (2012) measured halo masses from the COSMOS ACS data using a joint analysis of galaxy–galaxy weak lensing, galaxy spatial

clustering, and galaxy number densities. We find $\gamma = 0.047$ from this analysis. This value can be compared to previous studies that relate r_{50} to R_{vir} . These studies find $\gamma_{50} = 0.015\text{--}0.03$ (Kravtsov 2013; Huang et al. 2017; Jiang et al. 2018; Somerville et al. 2018). These values are consistent with our result for r_{80} , when the typical ratio between r_{80} and r_{50} is taken into account (a factor 2–3, depending on the Sérsic index).

The results are shown in Figure 3, and compared to SMHM relations from the literature. The derived SMHM function agrees very well with the lensing measurements from Leauthaud et al. (2012) at all masses and both redshift ranges where lensing data are available, even though we fit only for a single offset. Beyond $z \sim 1$ we cannot compare directly to measurements, but as shown in Leauthaud et al. (2012; and Figure 3) pivot halo mass measurements from lensing are consistent with those from halo occupation distribution (HOD) and subhalo abundance matching (SHAM) measurements. We therefore also include SMHM relations from SHAM and HOD

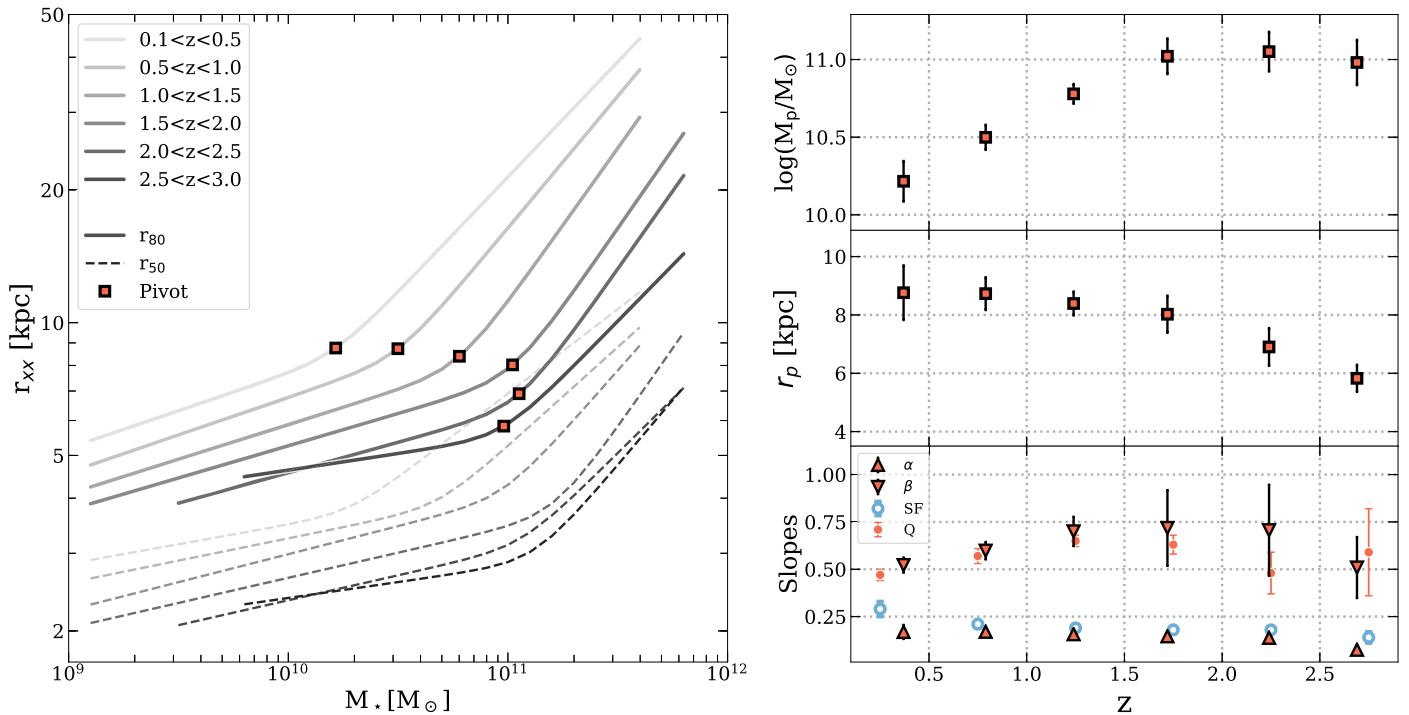


Figure 2. Redshift evolution of smoothly broken power-law fits (Equation (2)) to the size–mass relation. Left: evolution of size–mass relation of all galaxies since $z \sim 3$. The solid lines show the r_{80} –mass relation, while the broken lines show r_{50} –mass relation. The pivots of the broken power-law fits are indicated by red squares. Right: evolution of the parameters of broken power-law fits to median r_{80} –mass distribution. The top panel shows the pivot stellar mass, the middle panel shows pivot radius, and the bottom panel shows slopes α and β of the power law. Overplotted are the slopes of the single power-law fits to the star-forming and quiescent galaxies in Mowla et al. (2018).

measurements by Rodríguez-Puebla et al. (2017), Moster et al. (2013), Behroozi et al. (2018), and Legrand et al. (2018). At all redshifts the SMHM that we derive from galaxy sizes agrees well with that derived using other methods, although slope of the high-mass end and the pivot points start to diverge, as we will discuss later. This agreement with the much more sophisticated empirical modeling results is remarkable, especially given our simplistic assumption that, on average, $r_{80}/R_{\text{vir}} = 0.047$ across a wide range in stellar masses and cosmic epochs. In addition, the median halo mass at a given stellar mass is not necessarily equivalent to the median stellar mass at a given halo mass due to scatter in the relationship and the steepness of the mass function.

4.2. The Pivot Mass

A quantity that is of particular interest is the pivot mass; that is, the inflection mass where the slope of the SMHM changes. To compare the pivot stellar masses between various SMHMs, we fit SMHMs from the literature with our smoothly broken power-law relation (Equation (2)) using the same methodology as the fits to the r_{80} – M_* relation. The comparisons between the pivot stellar masses and pivot halo masses are shown in Figure 4. The pivot points are in good agreement, although we see a stronger evolution of pivot stellar mass in the size–mass relation than in SMHM relations from abundance matching up to $z \sim 3$. This has been noticed previously in Leauthaud et al. (2012), who found a more significant evolution of pivot stellar mass in SMHM measured from lensing and clustering between redshift 0.2 and 1 than from SHAM measurements.

5. Discussion

Using a new definition of size and the large galaxy sample from Mowla et al. (2018), we have shown that the size–mass relation of all (quiescent plus star-forming) galaxies is well fit with a broken power law. The stellar mass where the slope changes, the “pivot mass,” increases with redshift from $\log(M_p/M_\odot) \approx 10.2$ at $z = 0.25$ to $\log(M_p/M_\odot) \approx 11.0$ at $z = 2.75$. We have also shown that the form of this relation is remarkably similar to that of the SMHM. The pivot stellar masses of the two relations are identical within the errors, and the slope and normalization are very similar when the simple scaling $r_{80} = 0.047R_{\text{vir}}$ is assumed for all masses and redshifts. As discussed in Section 1, our results extend previous theoretical and observational studies (e.g., Kravtsov 2013; Huang et al. 2017, 2018; Somerville et al. 2018).

We note that our results do not rely on the use of r_{80} instead of r_{50} ; as shown in Figure 2, we derive similar relations for r_{50} , although the change in slope is not as striking as it is for r_{80} . The main advantages of r_{80} are that star-forming galaxies and quiescent galaxies have similar sizes at fixed stellar mass (see Miller et al. 2019) and that they encompass a larger fraction of the baryons.

It is interesting to speculate whether there is a straightforward physical interpretation of the similarity of these relations. From the stellar size–mass relation point of view, the pivot marks the stellar mass at which the galaxy population transitions from being dominated by star-forming galaxies to being dominated by quiescent galaxies. This is shown explicitly by the blue line in Figure 4, which indicates the mass where half the population is quiescent and half is star-forming. This evolving mass matches the pivot mass within the

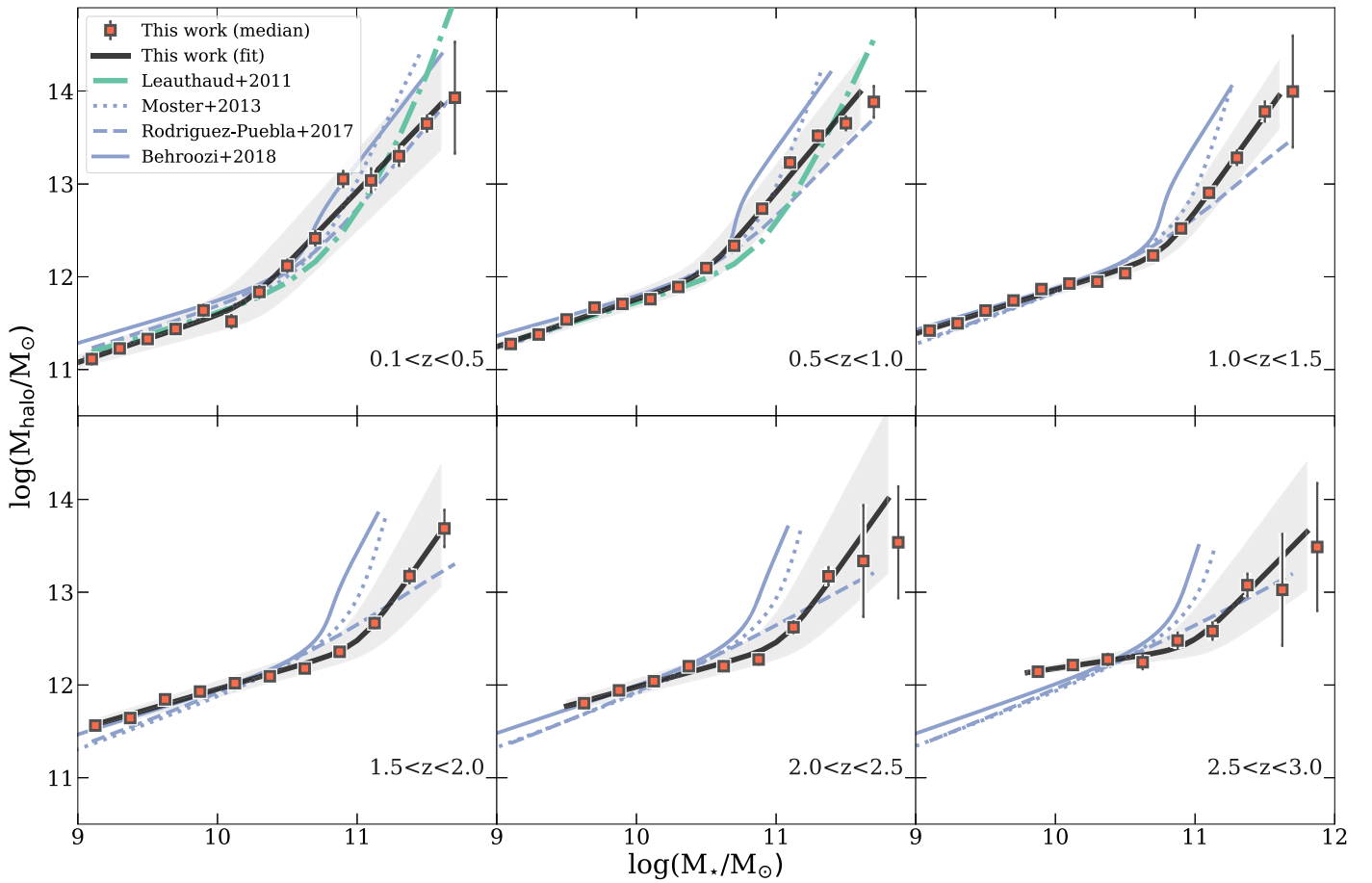


Figure 3. Halo mass–stellar mass relation derived from the size–mass relation of galaxies at $0.1 < z < 3.0$. The halo masses are determined from virial radii, with the simple assumption that the virial radius scales as $R_{\text{vir}} = \gamma^{-1} r_{80}$, with $\gamma = 0.047$. The red squares show the halo mass from median sizes, and the black line shows the halo mass–stellar mass relation from size–mass relation fit (gray band represents the error associated with the fit). Purple lines show stellar-to-halo mass relations derived using abundance matching techniques, and the green dashed line at $0.1 < z < 1.0$ is derived from galaxy–galaxy lensing and clustering (Leauthaud et al. 2012).

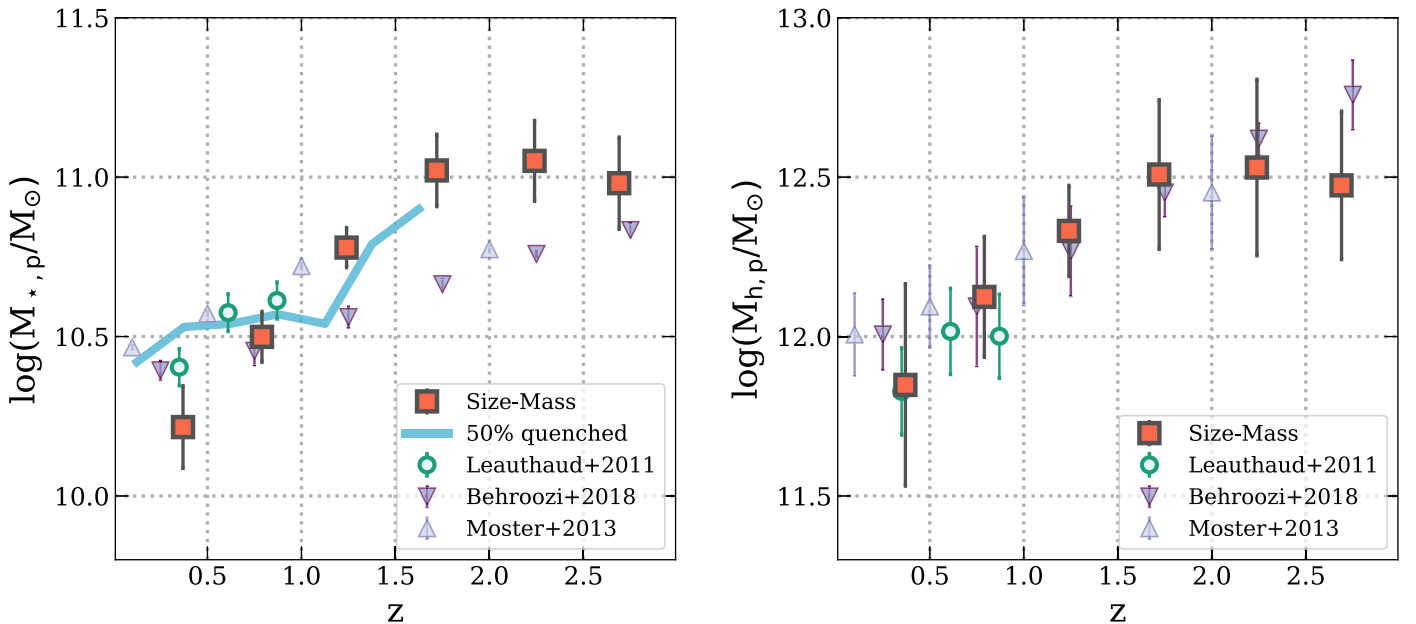


Figure 4. Redshift evolution of pivot points. The left panel shows the evolution of pivot stellar mass of the broken power-law fits to the size–mass relation (red circles), compared to similar fits to the stellar-to-halo mass relation from Behroozi et al. (2018), Leauthaud et al. (2012), and Moster et al. (2013). The blue line marks the stellar mass at which 50% of the galaxies are quenched (in the UVISTA catalog). The right panel shows the calculated halo mass at pivot point of size–mass relation compared to pivot halo masses stellar-to-halo mass relation.

errors, at least out to $z \sim 1.5$. From the stellar-to-halo mass relation point of view, the pivot is where M_*/M_h reaches a maximum; i.e., it is the halo mass at which baryons have been most efficiently converted into stars. Taking these aspects together, the pivot may simply mark the mass above which both the stellar mass growth and the size growth transition from being star formation dominated to being (dry) merger dominated (see also Dekel & Birnboim 2006).

It is not immediately obvious why the pivot mass should evolve with redshift. However, following Leauthaud et al. (2012), we note that the ratio of the pivot halo mass to the pivot stellar mass is roughly constant. That is, the evolution of the pivot mass and the size at the pivot mass conspire to keep the ratio M_{halo}/M_* approximately constant at the pivot mass (at ≈ 0.025).

Finally, we note that there are significant caveats associated with inverting an average stellar-to-halo mass relation to an average halo-to-stellar mass relation, as is done in Section 4. The existence of large, low surface brightness galaxies with very low velocity dispersions (Danieli et al. 2019), as well as the difference in clustering between star-forming and quiescent galaxies of the same stellar mass Coil et al. (2017), suggest that there is significant scatter in halo mass at fixed galaxy size. As discussed in detail by Somerville et al. (2018) the scatter in the stellar-to-halo mass relation, combined with the exponential fall-off in the stellar mass function, leads to an overestimate of the halo-to-stellar mass ratio at the high-mass end. Indeed, the inverted Moster et al. (2013) and Behroozi et al. (2018) relations are steeper than our derived relation. It is encouraging that our relation does agree with the direct estimates of Leauthaud et al. (2012), who measure average halo mass at fixed stellar mass without the need of conversions or assumptions about the scatter. Turning this argument around, the fact that we see a clear break in our inferred SMHM relation may imply a small scatter in the SMHM relation. We tested this by generating a mock catalog of galaxies using the SMHM from Behroozi et al. (2018) and introducing scatter in the stellar mass function. Beyond a scatter of 0.25 dex, the break in SMHM begins to disappear and the data can be reasonably well described by a single power law. In this framework our inferred relation implies a scatter of no more than 0.2 dex at the high-mass end, in line with other constraints (see Moster et al. 2013; Behroozi et al. 2018).

We thank Peter Behroozi, Andrey Kravtsov, Song Huang, and Johannes Lange for valuable feedback. We also thank the anonymous referee for insightful comments that improved the manuscript. This work is supported by NASA *HST* program GO-14114. A.v.d.W. acknowledges funding through the H2020 ERC Consolidator grant 683184.

ORCID iDs

Lamiya Mowla  <https://orcid.org/0000-0002-8530-9765>
 Arjen van der Wel  <https://orcid.org/0000-0002-5027-0135>
 Pieter van Dokkum  <https://orcid.org/0000-0002-8282-9888>

References

- Behroozi, P., Wechsler, R., Hearin, A., & Conroy, C. 2018, arXiv:1806.07893
 Behroozi, P. S., Conroy, C., & Wechsler, R. H. 2010, *ApJ*, 717, 379
 Bernardi, M., Meert, A., Vikram, V., et al. 2014, *MNRAS*, 443, 874
 Bryan, G. L., & Norman, M. L. 1998, *ApJ*, 495, 80
 Carollo, C. M., Bschorr, T. J., Renzini, A., et al. 2013, *ApJ*, 773, 112
 Coil, A. L., Mendez, A. J., Eisenstein, D. J., & Moustakas, J. 2017, *ApJ*, 838, 87
 Danieli, S., van Dokkum, P., Conroy, C., Abraham, R., & Romanowsky, A. J. 2019, arXiv:1901.03711
 Dekel, A., & Birnboim, Y. 2006, *MNRAS*, 368, 2
 Dutton, A. A., van den Bosch, F. C., Faber, S. M., et al. 2011, *MNRAS*, 410, 1660
 Elmegreen, D. M., Elmegreen, B. G., Ravindranath, S., & Coe, D. A. 2007, *ApJ*, 658, 763
 Ferguson, H. C., Dickinson, M., Giavalisco, M., et al. 2004, *ApJL*, 600, L107
 Genel, S., Nelson, D., Pillepich, A., et al. 2018, *MNRAS*, 474, 3976
 Huang, K.-H., Fall, S. M., Ferguson, H. C., et al. 2017, *ApJ*, 838, 6
 Huang, S., Leauthaud, A., Hearin, A., et al. 2018, arXiv:1811.01139
 Jiang, F., Dekel, A., Kneller, O., et al. 2018, arXiv:1804.07306
 Koekemoer, A. M., Aussel, H., Calzetti, D., et al. 2007, *ApJS*, 172, 196
 Koekemoer, A. M., Faber, S. M., Ferguson, H. C., et al. 2011, *ApJS*, 197, 36
 Kormendy, J. 1977, *ApJ*, 218, 333
 Kravtsov, A. V. 2013, *ApJL*, 764, L31
 Kravtsov, A. V., Vikhlinin, A. A., & Meshcheryakov, A. V. 2018, *AsL*, 44, 8
 Leauthaud, A., Tinker, J., Bundy, K., et al. 2012, *ApJ*, 744, 159
 Legrand, L., McCracken, H. J., Davidzon, I., et al. 2018, arXiv:1810.10557
 Massey, R., Stoughton, C., Leauthaud, A., et al. 2010, *MNRAS*, 401, 371
 Miller, T., van Dokkum, P., Mowla, L., et al. 2019, *ApJL*, 872, L14
 Mo, H. J., Mao, S., & White, S. D. M. 1998, *MNRAS*, 295, 319
 Mosleh, M., Williams, R. J., Franx, M., et al. 2012, *ApJL*, 756, L12
 Moster, B. P., Naab, T., & White, S. D. M. 2013, *MNRAS*, 428, 3121
 Mowla, L., van Dokkum, P., Brammer, G., et al. 2018, arXiv:1808.04379
 Muzzin, A., Marchesini, D., Stefanon, M., et al. 2013, *ApJS*, 206, 8
 Navarro, J. F., Benítez-Llambay, A., Fattahi, A., et al. 2017, *MNRAS*, 471, 1841
 Ono, Y., Ouchi, M., Curtis-Lake, E., et al. 2013, *ApJ*, 777, 155
 Peng, C. Y., Ho, L. C., Impy, C. D., & Rix, H.-W. 2010, *AJ*, 139, 2097
 Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016, *A&A*, 594, A13
 Rodríguez-Puebla, A., Primack, J. R., Avila-Reese, V., & Faber, S. M. 2017, *MNRAS*, 470, 651
 Shen, S., Mo, H. J., White, S. D. M., et al. 2003, *MNRAS*, 343, 978
 Skelton, R. E., Whitaker, K. E., Momcheva, I. G., et al. 2014, *ApJS*, 214, 24
 Somerville, R. S., Behroozi, P., Pandya, V., et al. 2018, *MNRAS*, 473, 2714
 Trujillo, I., Förster Schreiber, N. M., Rudnick, G., et al. 2006, *ApJ*, 650, 18
 van der Wel, A., Franx, M., van Dokkum, P. G., et al. 2014, *ApJ*, 788, 28
 van Dokkum, P. G., Nelson, E. J., Franx, M., et al. 2015, *ApJ*, 813, 23
 Williams, R. J., Quadri, R. F., Franx, M., et al. 2010, *ApJ*, 713, 738