

A Method of Correcting Near-Infrared Spectra for Telluric Absorption¹

WILLIAM D. VACCA,² MICHAEL C. CUSHING, AND JOHN T. RAYNER

Institute for Astronomy, University of Hawaii, Honolulu, HI 96822; vacca@mpe.mpg.de, cushing@ifa.hawaii.edu, rayner@ifa.hawaii.edu

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ABSTRACT. We present a method for correcting near-infrared medium-resolution spectra for telluric absorption. The method makes use of a spectrum of an A0 V star, observed near in time and close in air mass to the target object, and a high-resolution model of Vega, to construct a telluric correction spectrum that is free of stellar absorption features. The technique was designed specifically to perform telluric corrections on spectra obtained with SpeX, a 0.8–5.5 μm medium-resolution cross-dispersed spectrograph at the NASA Infrared Telescope Facility, and uses the fact that for medium resolutions there exist spectral regions uncontaminated by atmospheric absorption lines. However, it is also applicable (in a somewhat modified form) to spectra obtained with other near-infrared spectrographs. An IDL-based code that carries out the procedures is available for downloading via the World Wide Web.

1. INTRODUCTION

Ground-based near-infrared spectroscopy ($\sim 1\text{--}5 \mu\text{m}$) has always been hampered by the strong and variable absorption features due to the Earth's atmosphere. Even within the well-established photometric bands such as *J*, *H*, and *K*, so-called telluric absorption features are present. The common method used to correct for such absorption features, and to retrieve the signal of the object under observation, is to observe a "telluric standard star," near in both time and sky position (air mass) to the object, and subsequently to divide the object's spectrum by the standard star's spectrum during the final phases of the reduction process.

Early A-type stars are frequently used as such telluric standards, as their spectra contain relatively few metal lines, with strengths that are generally only a few percent of the continuum. Furthermore, for A stars with little reddening, the magnitudes at all wavelengths can be deduced from a single (usually *V*-band) measurement. In addition, the near-infrared continua of A stars can be reasonably well approximated by a blackbody with a temperature of $\sim 10,000$ K. These latter two properties make A-type stars very useful for determining the absolute flux calibration of the observations. However, the strong intrinsic photospheric hydrogen (H) absorption features in the spectra of A-type stars present a substantial problem for their use as telluric standards, particularly if one is working in the *I*, *z*, *Y*, or *H* bands (where the numerous higher level lines of the Paschen and Brackett series are located), or if the observations

are designed to study the H lines in the spectrum of the object of interest. One technique commonly adopted in order to preserve the other desirable qualities of A stars as telluric and flux standards is to simply interpolate over the H absorption lines in their observed spectra. However, since atmospheric absorption features are certainly present at the locations of most of these lines, this method can lead to uncertainties in the reality of any features subsequently detected in the corrected object's spectrum at or near the wavelengths of the stellar H lines. Furthermore, A-type stars can have rather large rotational velocities (upwards of 300 km s^{-1} ; e.g., Abt & Morrell 1995; Royer et al. 2002), which broaden the absorption lines. The large number of such broad lines renders the interpolation method infeasible near the various H series limits.

Maiolino, Rieke, & Rieke (1996) developed a correction method using G2 V stars, rather than A-type stars, as telluric standards. Their method makes use of the high signal-to-noise ratio (S/N), high-resolution spectrum of the Sun, obtained using a differential air mass technique at the McMath Telescope at Kitt Peak National Observatory by Livingston & Wallace (1991; see also Wallace, Hinkle, & Livingston 1993), to generate a telluric correction spectrum, which is then applied to the spectrum of the object of interest. The correction spectrum is derived from the ratio of the solar spectrum (shifted to the appropriate radial velocity and convolved to the resolution of the observations) to the observed G2 V star spectrum. This ratio should, in principle, remove any intrinsic stellar features from the telluric correction spectrum. Although Maiolino et al. (1996) demonstrated that this method can be reasonably successful in performing telluric corrections on moderate S/N and resolution data, there are several practical problems with it, especially when applying it to higher quality and resolution spectra. First, the number of reasonably bright G2 V stars avail-

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² Currently at the Max-Planck-Institut für Extraterrestrische Physik, Gies-senbachstrasse, 85741 Garching, Germany.

TABLE 1
A0 V AND G2 V STARS IN
SIMBAD DATABASE

V (mag)	$N(\text{A0 V})$	$N(\text{G2 V})$
<7	364	50
<8	740	171
<9	1687	459
<10	3633	993
<11	4557	1165

NOTE.—Cumulative number of A0 V and G2 V stars listed in the SIMBAD database as of 2000 July, as a function of V -band magnitude.

able for use as telluric standards near any given object is less than the number of A-type stars, simply because G stars are inherently less luminous than A stars and catalogs of stars with well-defined spectral classifications are based on magnitude-limited samples. This is demonstrated in Table 1, in which we have compiled the numbers of stars listed with A0 V and G2 V spectral types in the SIMBAD database as a function of V magnitude. The intrinsic ($V-K$) color of a G2 V star is about 1.5 mag; therefore, even in the K band, bright A0 V stars outnumber G2 V stars in the database. Second, G-type stars have a large number of (relatively weak) metal lines throughout their spectra. Hence, the accuracy of the resulting corrections depends on how well the observed spectrum of the G star matches that of the Sun; the higher the S/N and resolution of the data, the closer the spectral match must be for a desired level of accuracy. As the strengths of the metal lines vary considerably with spectral type and stellar parameters, any mismatch, such as might be caused by spectral classification inaccuracies, metal abundance differences, anomalous line strengths, etc., will introduce spurious features into the telluric correction spectrum. Third, the reference solar spectrum was obtained at a relatively low site and has large wavelength gaps where the intrinsic spectrum could not be recovered using the differential air mass method. Because the depths of atmospheric absorption features vary with elevation, this solar spectrum would not provide a useful correction template for spectra, such as those obtained at a high site such as Mauna Kea, which have recoverable flux in the regions of strong absorption between the photometric passbands.

An alternative method, used, for example, by Hanson, Conti, & Rieke (1996), combines the A-star technique with that of Maiolino et al. (1996) and involves the following steps: observe both an A-type star and a G2 V star, normalize the spectra of both stars, “correct” the normalized G2 V star spectrum by dividing it by the solar spectrum, replace the regions centered on the H lines in the normalized A-star spectrum by the corrected G-star spectrum, and finally divide the result into the object’s spectrum. This is equivalent to dividing the object spectrum by that of the A star and multiplying the result by a line-correction spectrum generated from the G2 V star and the

solar spectrum. Since the solar spectrum is used only in the regions of the H lines, the resulting correction template should be more accurate than one based on the G2 V star/solar spectrum ratio alone. Because the normalization process removes the signature of the instrumental throughput from the observed stellar spectra, the object must be flux-calibrated separately, after the correction process, usually by multiplying the normalized, corrected spectrum by a blackbody with a temperature appropriate for an A star. While this method might seem attractive in principle, in practice it requires observations of two sets of standards (a suitable one of which may be difficult to find, as explained above) for each object of interest (ideally) and involves a somewhat tedious and complicated reduction procedure, particularly if the wavelength range of interest includes several absorption lines of the Pa or Br series.

Clearly, the ideal telluric standard would be one whose spectrum is completely featureless and precisely known. In addition, such objects should be reasonably bright, fairly common, and widely distributed on the sky, so one can be found close to any particular object of interest. While few objects satisfying the first criterion are known in nature, there are a limited number of objects for which precise and accurate theoretical models are available. Vega is one such object. In this paper, we describe a new method of generating a telluric correction spectrum using observations of A0 V stars and a high-resolution model of Vega’s intrinsic spectrum. Through a process somewhat similar to that described by Maiolino et al. (1996), this technique preserves the desirable qualities of A-type stars, while avoiding the problems with the intrinsic stellar absorption lines, and does not require observations of G2 V stars.

The method outlined below was developed for the reduction of spectra obtained with SpeX at the NASA Infrared Telescope Facility (IRTF) on Mauna Kea. SpeX is a medium-resolution ($\lambda/\Delta\lambda \sim 200\text{--}2500$) cross-dispersed spectrograph equipped with a 1024×1024 InSb array that provides simultaneous wavelength coverage over $0.8\text{--}2.4 \mu\text{m}$ in one grating setting and $2.4\text{--}5.5 \mu\text{m}$ in another grating setting (see Rayner et al. 1998, 2003 for a full description of SpeX and its various scientific observing modes). Because the entire Paschen series and most of the Brackett series of H absorption lines are simultaneously detected in a spectrum of any A0 V star observed with SpeX, it was important to devise a technique that could be applied easily and rapidly to accurately remove all of these features from a telluric correction spectrum generated from the A0 V star spectrum. With high-quality spectra of an A0 V star, a telluric correction spectrum accurate to better than about 2% across the H lines can be quickly constructed with this technique. Although designed specifically for SpeX data, the method can also be used to correct spectra obtained with most-medium resolution ($\lambda/\Delta\lambda < 50,000$) near-infrared spectrographs. A set of IDL-based routines (Spextool_extension, version 2.0), specifically designed for application to SpeX data, implements the telluric correction method described below and forms an extension to Spextool (version 2.3), a package of

routines written for the basic reduction of spectra obtained with SpeX (M. C. Cushing, W. D. Vacca, & J. T. Rayner 2003, in preparation). Both sets of programs are available from the SpeX Web page at the IRTF Web site.³ Finally, a third set of IDL routines, constituting a general version (i.e., not specific to SpeX data) of the telluric correction code, is also available from this Web page.

2. DESCRIPTION OF METHOD

We assume that the observed spectrum $O(\lambda)$ of any object is given by

$$O(\lambda) = [I(\lambda) \cdot T(\lambda)] * P(\lambda) \cdot Q(\lambda), \quad (1)$$

where $I(\lambda)$ is the intrinsic spectrum of the object, $T(\lambda)$ is the atmospheric (telluric) absorption spectrum, $P(\lambda)$ is the instrumental profile, $Q(\lambda)$ is the instrumental throughput, and the asterisk denotes a convolution. Because the convolution represents a smoothing of the intrinsic stellar and telluric spectra, the above equation can be rewritten as

$$O(\lambda) = [I(\lambda) * P(\lambda)] \cdot [T(\lambda) * P(\lambda)] \cdot Q(\lambda) \quad (2)$$

or

$$O(\lambda) = I_{\text{instr}}(\lambda) \cdot R(\lambda), \quad (3)$$

where

$$I_{\text{instr}}(\lambda) = I(\lambda) * P(\lambda) \quad (4)$$

is the spectrum of the object observed above the atmosphere with an instrument with perfect throughput and

$$R(\lambda) = [T(\lambda) * P(\lambda)] \cdot Q(\lambda) \quad (5)$$

is the telluric absorption spectrum convolved with the instrument profile and multiplied by the instrumental throughput. The latter represents the overall throughput, or response, of the atmosphere+instrument system. We assume that $Q(\lambda)$ is a smooth, featureless function and therefore any observed high-frequency variations in $O(\lambda)$ must be due to either $I(\lambda)$ or $T(\lambda)$.

The immediate objective is to determine I_{instr} of a specific object from $O(\lambda)/R(\lambda)$. Conversely, if $I(\lambda)$ for any particular source is known a priori, and $P(\lambda)$ of the instrument is known or can be determined, then the system throughput $R(\lambda)$ can be determined from equations (3) and (4). This factor can be subsequently used to determine $I_{\text{instr}}(\lambda)$ of another (target) object from the observed spectrum, provided that the telluric absorption spectrum $T(\lambda)$ does not vary substantially between the two sets of observations. This forms the basis of both our method and that of Maiolino et al. (1996): observations of a “telluric

standard,” whose intrinsic spectrum is known, to derive a system throughput curve that can then be applied to a target object spectrum observed nearby in the sky and close in time. Since $T(\lambda)$ varies with air mass and on timescales of the order of several to tens of minutes (depending on the atmospheric conditions), it is clearly best to observe a telluric standard at an air mass as close as possible to that of the target object and within a few minutes of the observations of the target.

Because the metal lines in A0 V spectra are relatively infrequent and weak (on the order of a few percent of the continuum level in the near-infrared at the resolution of SpeX; see, for example, Fig. 1), the intrinsic near-infrared spectra of such objects consist essentially of H absorption lines and a smooth, featureless (blackbody) continuum. Thus, metal abundance variations do not strongly affect the appearance of the observed spectra and, except for possible differences in the widths and depths of some of the H lines (due to different rotational velocities or slightly different values of the surface gravity), the spectra of all A0 V stars can be considered nearly identical to one another. It is a fortunate circumstance that Vega, the object that defines the zero point of the magnitude system, is also a bright A0 V star whose spectrum has been accurately modeled. Assuming that the model provides an accurate representation of the intrinsic spectrum of Vega and that Vega is a good representative (an archetypal example) of the class of A0 V stars, we can use the theoretical model to determine $I_{\text{instr}}(\lambda)$ of any A0 V star. Hence, if each observation of a given target object is preceded and/or followed by an observation of an A0 V star, and if the model spectrum of Vega can be scaled and reddened to match the near-infrared magnitudes of the observed A0 V star and modified to account for the differences in line strengths, radial and rotational velocities, and spectral resolution, we will be in the happy situation of having at hand both the observed spectrum $O(\lambda)$ and intrinsic spectra $I(\lambda)$ and $I_{\text{instr}}(\lambda)$ of the A0 V star. The system response function $R(\lambda)$ can then be determined immediately and used to correct the target object’s spectrum. The difficulty then lies “only” in properly modifying the model spectrum of Vega.

Modifying the model Vega spectrum involves shifting it to the radial velocity of the observed A0 V star, scaling and reddening it to reproduce the known magnitudes of the observed A0 V star, convolving it with a function (kernel) that broadens the lines to the observed widths and smooths it to the observed resolution, resampling the result to the observed wavelength scale, and altering the depths of various H lines to match those of the observed A0 V star. The convolution kernel is given by the instrument profile (IP) $P(\lambda)$ and a (rotational) broadening function $\Theta(\lambda)$ (see below). There are at least three ways of constructing $P(\lambda)$: (1) assume a theoretical profile whose properties are defined by the instrument parameters (e.g., slit width), (2) construct an empirical profile using observations of unresolved arc lines (similar to those used to perform the wavelength calibration of the spectra), and (3) construct a profile using the observations of the A0 V star itself. We discuss the

³ <http://irtfweb.ifa.hawaii.edu/Facility/spex>.

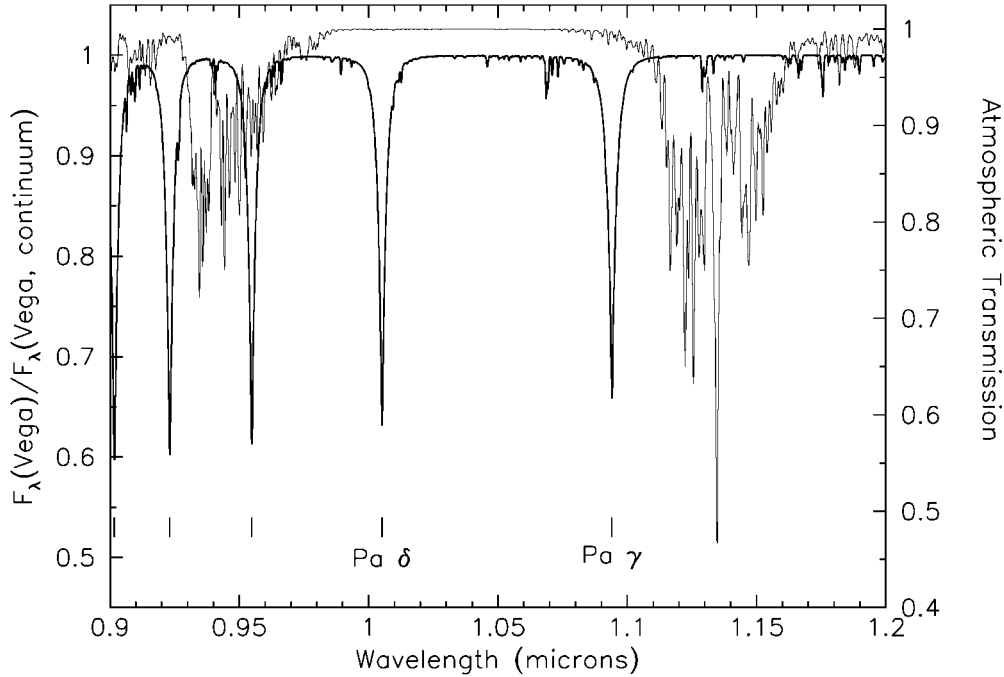


FIG. 1.—Normalized model spectrum of Vega, smoothed to a resolving power of 2000 (*thick line, left-hand axis*). The theoretical atmospheric transmission around the Pa δ line (1.00494 μm) at air mass 1.2 for typical conditions on Mauna Kea, smoothed to a resolving power of 2000 (*thin line, right-hand axis*). Note that the Pa δ line is located in a region where atmospheric absorption is negligible.

last method in detail here and describe below a set of routines that implement this method for use in the reduction of data obtained with SpeX. (The routines also allow users to adopt empirical or theoretical instrument profiles.)

We begin by selecting a wavelength region centered on an intrinsic absorption line in the observed and model A0 V spectra. The observed spectrum is normalized in the region of the line by fitting a low-order polynomial or a spline to the continuum; a good representation of the observed continuum can be obtained once bad pixels and the regions around any strong absorption features are excluded from the fitting process. The normalization removes the signature of the instrumental throughput as well as the shape of the stellar continuum. Once the observed stellar spectrum has been normalized, the wavelength region around the absorption feature can be cross-correlated with that from the normalized model of Vega, using a technique similar to that described by Tonry & Davis (1979), to determine the observed radial velocity of the A0 V star. The full (i.e., unnormalized) model is then shifted by the estimated radial velocity and scaled and reddened to match the observed magnitudes of the A0 V star.

We now wish to construct a convolution kernel that, when applied to the model spectrum of Vega, will produce a close approximation of the intrinsic spectrum of the observed A0 V star, as measured by an instrument with perfect throughput. That is, we assume that the intrinsic spectrum of the A0 V star

is given by

$$I_{\text{A0 V}}(\lambda) = I_{\text{Vega}}(\lambda) * \Theta(\lambda), \quad (6)$$

where $I_{\text{Vega}}(\lambda)$ is the scaled and shifted model spectrum of Vega and $\Theta(\lambda)$ is a function that accounts for the differences in rotational velocity and any other line-broadening factors between the Vega model and the observed A0 V star. The observed A0 V spectrum $O_{\text{A0 V}}(\lambda)$ is then given by

$$O_{\text{A0 V}}(\lambda) = [I_{\text{Vega}}(\lambda) * \Theta(\lambda)] \cdot T(\lambda) * P(\lambda) \cdot Q(\lambda) \quad (7)$$

or

$$O_{\text{A0 V}}(\lambda) = [I_{\text{Vega}}(\lambda) * K(\lambda)] \cdot R(\lambda), \quad (8)$$

where

$$K(\lambda) = \Theta(\lambda) * P(\lambda), \quad (9)$$

$$I_{\text{instr,A0 V}} = I_{\text{Vega}}(\lambda) * K(\lambda). \quad (10)$$

For the moment, let us simply assume we can find an intrinsic stellar absorption line in the observed A0 V spectrum across which both the atmospheric absorption and the instrumental throughput are either constant or only smoothly varying. Using the normalized observed and model spectra in the vicinity of

the line, we then have

$$O'_{A0V}(\lambda) = I'_{Vega}(\lambda) * K(\lambda), \quad (11)$$

where primes denote the normalized spectra (divided by the respective continua). The convolution kernel $K(\lambda)$ can then be determined by selecting a small region around the observed absorption line and performing a deconvolution.

Once $K(\lambda)$ has been determined, it is convolved with the shifted, scaled, and reddened Vega model spectrum and $R(\lambda)$ is determined from equation (8). Under the assumptions that all A0 V stars are identical to Vega and that the model of Vega provides a highly accurate representation of Vega's intrinsic spectrum, the convolution and subsequent division $O_{A0V}(\lambda)/[I_{Vega}(\lambda) * K(\lambda)]$ should yield a response spectrum in which the intrinsic stellar absorption lines are absent. In practice, we find that some small residual features at the locations of the H lines are discernible in this response spectrum because of differences in the strengths of these lines between the model and the observed A0 V spectra. Some of these differences may in fact be real, due to different values of $\log g$ between the standard star and the Vega or misclassification of the observed standard star and hence a slight spectral mismatch between the standard and Vega. However, the model H line strengths (equivalent widths) can be easily adjusted to minimize, or eliminate, these residuals. Furthermore, since any metal lines are fairly weak and do not vary strongly among late B-type and early A-type stars, the effects of spectral mismatches and misclassifications on the resulting telluric correction spectrum can be minimized to a large extent by simply modifying the H line strengths in the model. For the resolving power and typical S/N levels achieved with SpeX, minor misclassifications should have a negligible effect on the final corrected spectra.

The intrinsic spectrum of the target object $I_{instr}(\lambda)$ is then determined from equation (3), rewritten as

$$I_{instr}(\lambda) = O(\lambda) \cdot [I_{Vega}(\lambda) * K(\lambda)/O_{A0V}(\lambda)]. \quad (12)$$

We refer to the quantity $[I_{Vega}(\lambda) * K(\lambda)]/O_{A0V}(\lambda) = 1/R(\lambda)$ as the “telluric correction spectrum.” One advantage of this method is that the resulting spectrum $I_{instr}(\lambda)$ is automatically placed on the proper flux scale if the slit losses during the two sets of observations (standard star and target object) were the same.

In summary, the method of constructing a telluric correction spectrum consists of an observation of an A0 V star and several steps designed to modify the model spectrum of Vega appropriately:

1. normalization of the observed A0 V star spectrum in the vicinity of a suitable absorption feature (as defined below);
2. determination of the radial velocity shift of the A0 V star;
3. shifting the Vega model spectrum to the radial velocity of the A0 V star;

4. scaling and reddening the Vega model spectrum to match the observed magnitudes of the A0 V star;

5. construction of a convolution kernel from a small region around an absorption feature in the normalized observed A0 V and model Vega spectra;

6. convolution of the kernel with the shifted, scaled, and reddened model of Vega;

7. scaling the equivalent widths of the various H lines to match those of the observed A0 V star.

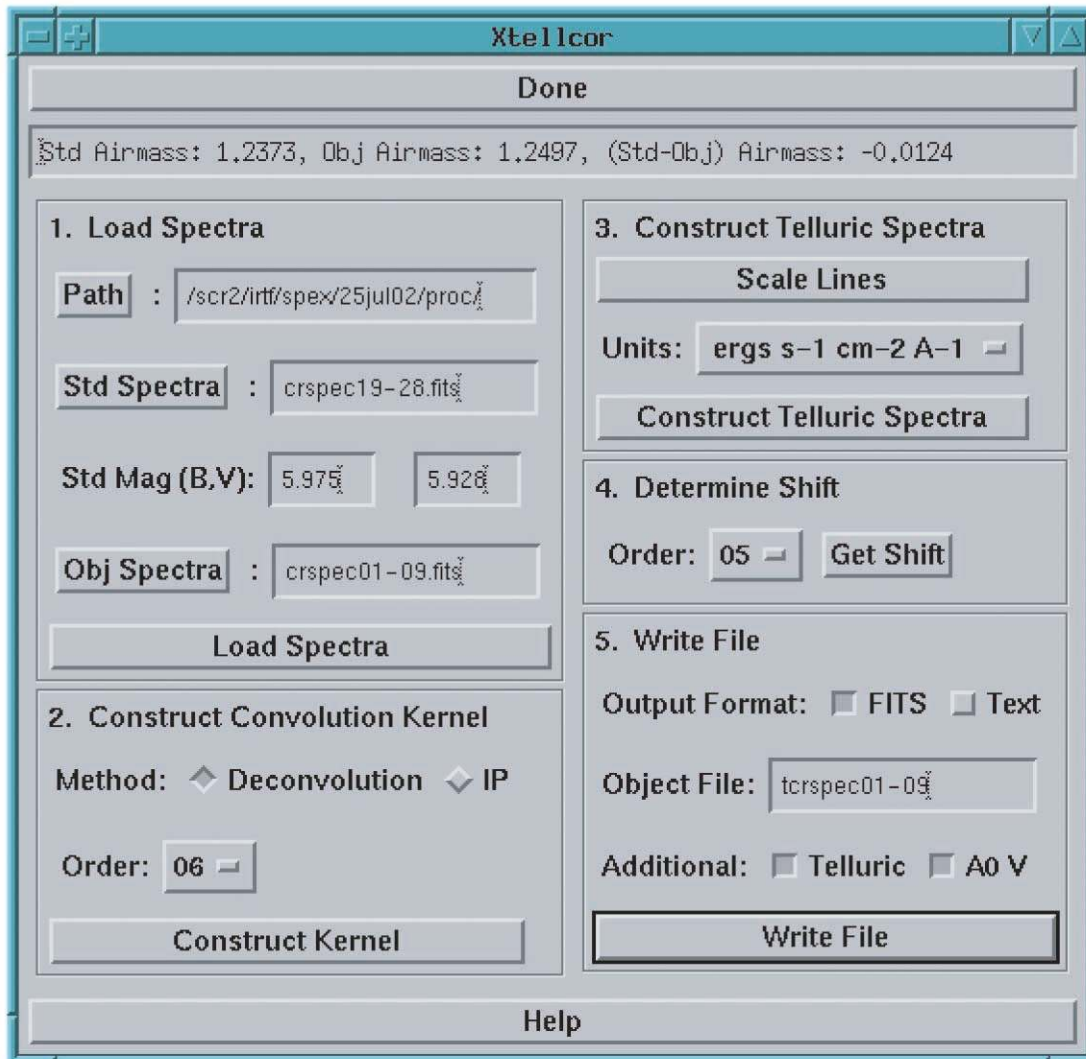
Finally, the convolved model is divided by the observed A0 V spectrum and the resulting telluric correction spectrum is multiplied by the observed target spectrum.

3. PRACTICAL CONSIDERATIONS

For the method described above to work properly, the convolution kernel $K(\lambda)$ must be determined accurately. If $P(\lambda)$ and $\Theta(\lambda)$ are known in advance, $K(\lambda)$ can be constructed using equation (9). If $K(\lambda)$ is to be generated from the observations themselves, it is necessary to find a region of the observed A0 V spectrum that contains a strong intrinsic stellar absorption line (so that the radial and rotational velocities and IP can be determined) but is free from atmospheric absorption or instrumental features. The latter requirement arises from our assumption that $T(\lambda)$ and $Q(\lambda)$ either are constant or vary only slowly and smoothly across the width of the line. Fortunately, for observations with SpeX, the atmosphere has cooperated in providing such a wavelength region and stellar atmospheres have cooperated in providing a suitable absorption line. Figure 1 presents a model of the atmospheric transmission between 0.9 and 1.2 μm , computed with the ATRAN program (Lord 1992) for an air mass of 1.2 and standard conditions on Mauna Kea and smoothed to a resolving power of 2000. The normalized model spectrum of Vega obtained from R. Kurucz,⁴ also smoothed to resolving power of 2000, is overplotted. It can be seen that the Pa δ (1.00494 μm) absorption line in the Vega spectrum is located in a region where, at this resolution, atmospheric absorption is almost completely negligible. Therefore, this line can be used to determine the convolution kernel $K(\lambda)$ in the procedure described above.

In addition to the lack of telluric contamination, the Pa δ line also lies in a wavelength region where the throughput $Q(\lambda)$ of SpeX is relatively high and does not change dramatically across the width of the line. The throughput around the Pa γ (1.0938 μm) line, for example, drops quickly across the line (due to the grating blaze function) and atmospheric absorption affects the shape of the long-wavelength wing, making the determination of the continuum near the wings difficult. Nevertheless, although the Pa δ line is ideal for our purposes, we have successfully generated telluric correction spectra using the method described in § 2 with kernels built from other H lines in the spectra of A0 V stars (e.g., Br γ). For spectrographs that do not

⁴ <http://kurucz.harvard.edu/stars.html>.

FIG. 2.—Control panel widget for *xtellcor*.

contain cross-dispersers and/or yield only a limited wavelength range, the observed A0 V spectra may not contain a line that is suitable for constructing $K(\lambda)$. In these cases, the convolution kernel can be generated from the IP and some estimate of the rotational velocity of the A0 V star (eq. [9]). At resolving powers similar to those of SpeX, the latter generally has a negligible effect on the resulting telluric correction spectrum.

The model of Vega we use for our telluric correction procedure has a resolving power ($\lambda/\Delta\lambda$) of 500,000, binned by five (giving an effective resolving power of 100,000), and has been scaled to match the observed Vega flux at 5556 Å given by Mégessier (1995). (A higher resolution model is also available from the Kurucz Web page; however, the array lengths necessary to store these data and the subsequent intermediate products become inordinately large and unwieldy.) The properties of the model are as follows: $T_{\text{eff}} = 9550$ K, $\log g =$

3.950, $v_{\text{rot}} = 25$ km s $^{-1}$, $v_{\text{turb}} = 2$ km s $^{-1}$, and a metal abundance somewhat below solar. The model also has an associated continuum spectrum; using the full spectrum and the continuum, we generated a normalized line spectrum. The continuum spectrum contains sharp jumps due to the abrupt increases/decreases in the opacity at the limits of the H (Pa and Br) series. The strengths of these jumps must be modified along with the line strengths. To account for this, we have normalized the continuum spectrum by dividing it by a spline fitted to the spectrum far from the jumps. The fit passes approximately through the flux midpoints between the continuum levels on either side of the jumps. Scaling this normalized continuum by an appropriate value then has the effect of scaling the heights of the jumps.

Note that because the Vega model spectrum has such a large resolving power, the method we have outlined should work

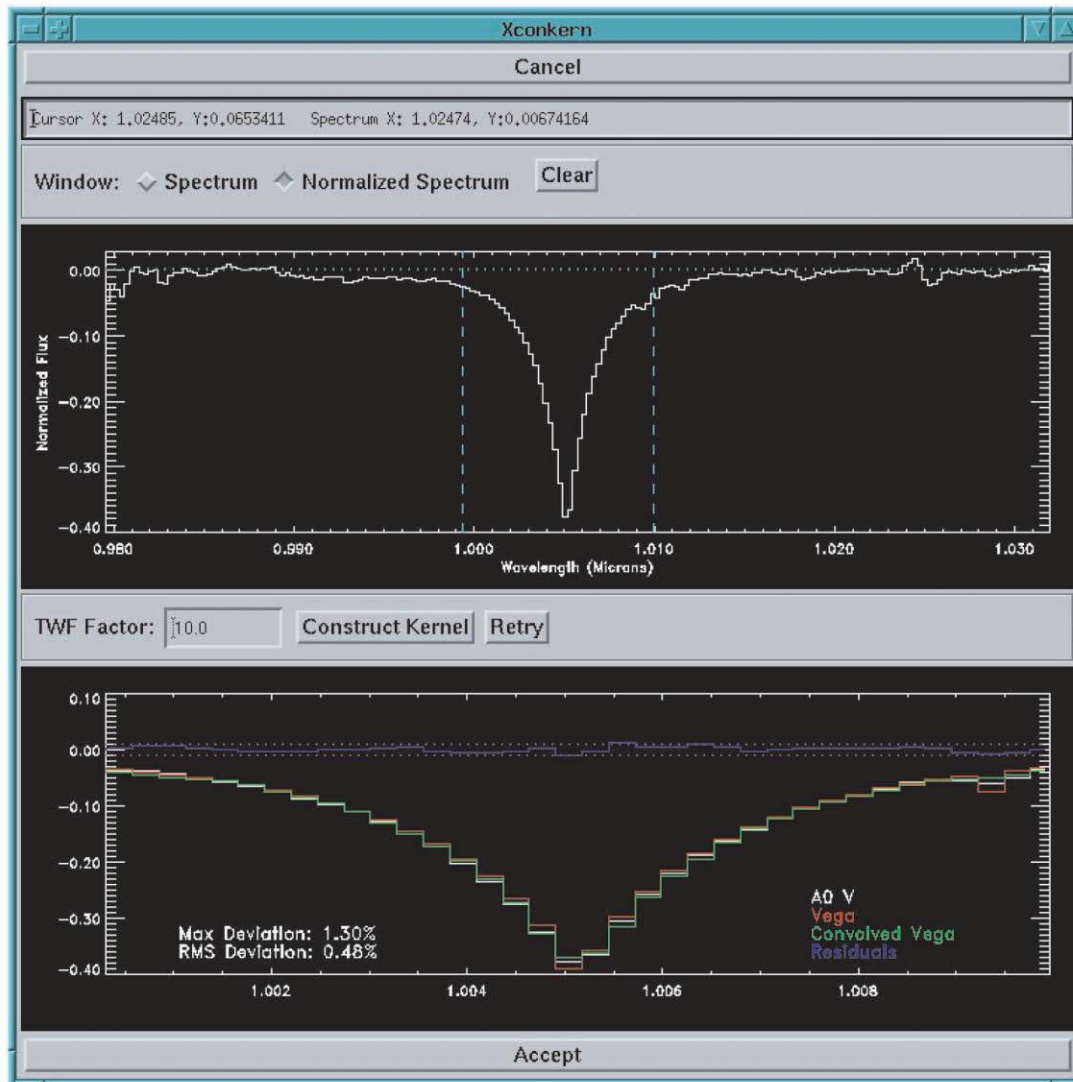


FIG. 3.—Determination of the convolution kernel from the Pa δ line in the normalized spectrum of an A0 V star. The vertical dashed lines denote the region of the spectrum used in the deconvolution to generate the kernel. Note that the residuals (seen in the bottom panel) have an rms deviation of much less than 1%.

successfully for spectral data with considerably larger resolving powers than that of SpeX ($\lambda/\Delta\lambda = 200\text{--}2500$). In principle, use of the routines described below (§ 4) is limited to data with $\lambda/\Delta\lambda < 50,000$, although data with higher resolving powers could be corrected if a model with higher resolving power were incorporated into the code. We have used a more general version of the code described below to successfully perform telluric corrections on data with resolving power of 5000, and we routinely use the code, with an empirical instrument profile $P(\lambda)$, to successfully correct prism data ($\lambda/\Delta\lambda \sim 200$) obtained with SpeX. At low resolving powers, the possibility of finding isolated H lines that are uncontaminated by atmospheric absorption is greatly diminished and precludes the construction of the convolution kernel $K(\lambda)$ from the observed data. At

$\lambda/\Delta\lambda \sim 200$, for example, even the Pa δ line is affected by telluric absorption in the wings.

The calculations described above (§ 2) are easiest to carry out in velocity coordinates, which is appropriate for a spectrum with constant spectral resolving power. In addition, the Vega model has a constant resolving power. However, in most spectrograph designs, the resolving power varies across the spectral range. SpeX is no exception, and the resolving power varies linearly by about 20% across each order (Rayner et al. 2003). Fortunately, the spectral resolution $\Delta\lambda$ of SpeX is constant to a high degree across each order, although the specific value varies from order to order. However, in pixel coordinates it is the same value for all orders, given approximately by the projected slit width. This implies that, in pixel coordinates, a kernel

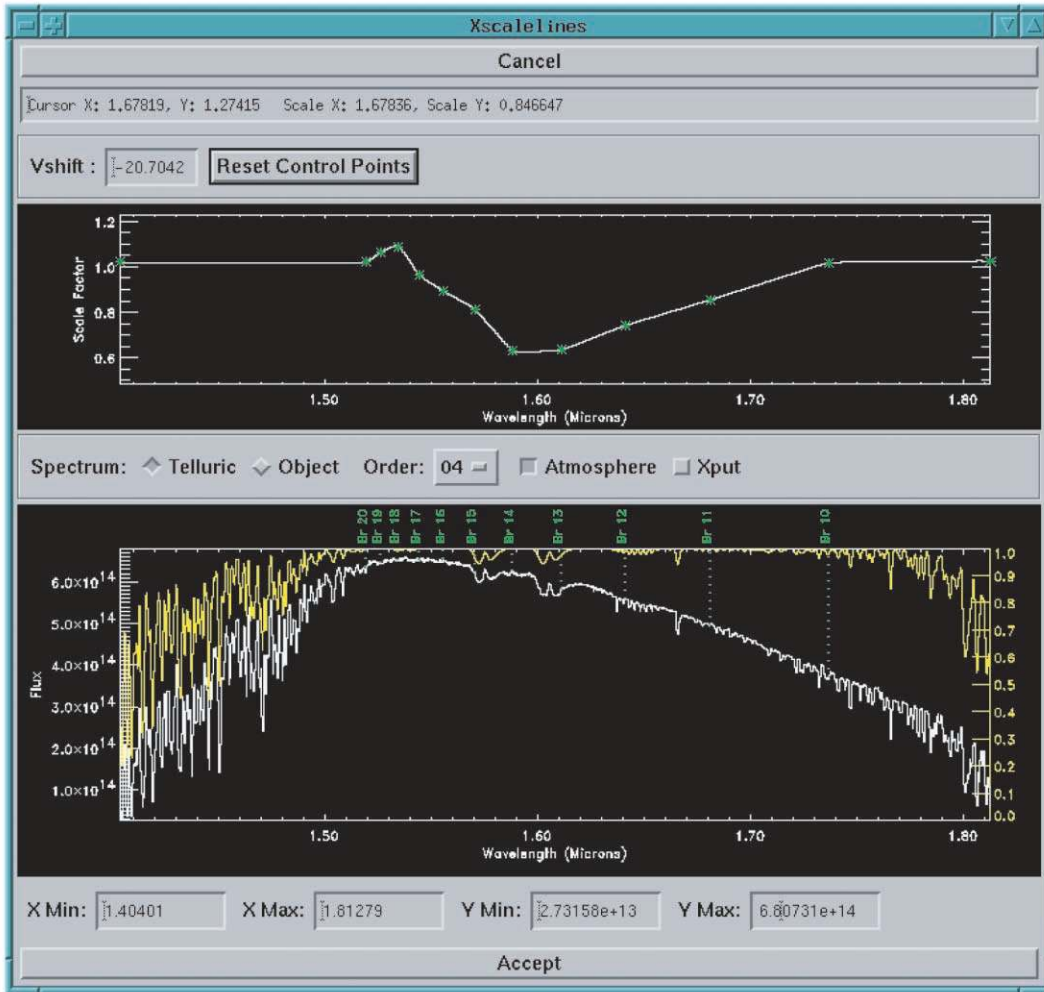


FIG. 4.—Determination of the H line scale factors for the Vega model in order 4 of SpeX. The upper panel shows the scale factors for each of the H absorption lines, while the bottom panel shows the response spectrum. The yellow curve is the theoretical atmospheric transmission for an air mass of 1.2 and standard conditions on Mauna Kea.

$K(\lambda)$ built from a line in one order can be applied to all other orders. Therefore, the code described below carries out the deconvolution in wavelength coordinates in order to construct the kernel, converts the resulting $K(\lambda)$ to pixel coordinates, and then applies this kernel to each order. This saves the user from the tedious task of building kernels specific to each order.

Deconvolution procedures are notoriously sensitive to noise. To minimize the effects of noise on the deconvolution needed to construct $K(\lambda)$, we multiply the ratio of the Fourier transforms of $O'_{A0V}(\lambda)$ and $I'_{Vega}(\lambda)$ by a tapered window function of the form $1/[1 + (\nu/\nu_{\max})^{10}]$, where ν_{\max} is a reference frequency given by a multiple of the width of a Gaussian fit to the Fourier transforms. This has the effect of suppressing high-frequency noise.

Clearly, as the response spectrum $R(\lambda)$ will be divided into the observed target spectrum, it is also important to acquire high S/N values in the observed spectrum of the A0 V star in

order to achieve accurate and reliable telluric corrections. In general, we have found that an S/N of ≥ 100 is needed in order to obtain telluric-corrected spectra that are limited by the flat-fielding errors and the S/N of the observed target spectrum. Of course, the accuracy of the final corrected spectra also depends on the stability of the atmospheric conditions, the proximity (in both time and air mass) of the A0 V observations to those of the target object, and the degree to which the A0 V spectrum matches the model spectrum of Vega. However, as we have argued above, small spectral mismatches should have a minor effect on the generated telluric correction spectrum. Fortunately, S/N values of 100 are rather easy to achieve at resolving powers of ≤ 2500 on 3 m class telescopes. Using SIMBAD, we have compiled a list of A0 V stars that are suitable for carrying out the method described above. The SpeX Web page includes a Web form that will generate a listing of the known A0 V stars within a specified angular distance

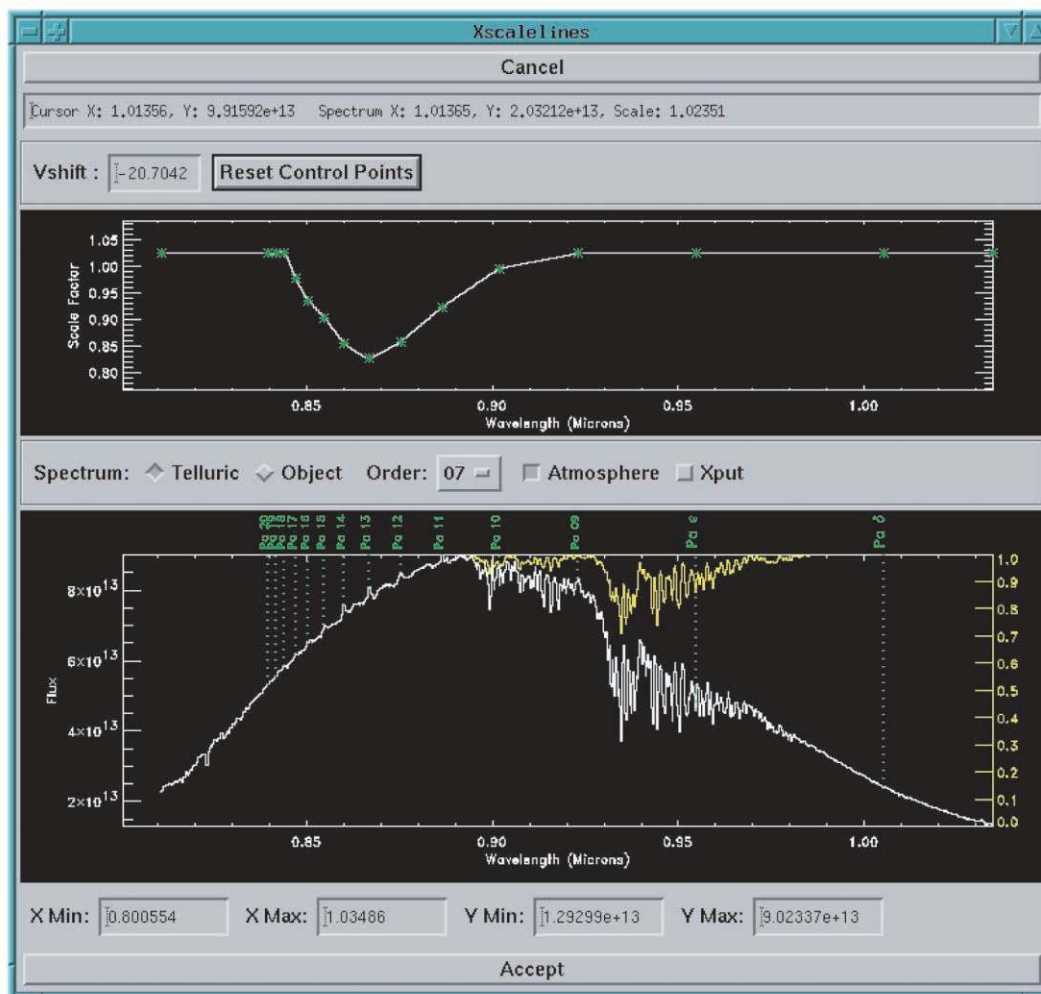


FIG. 5.—Same as Fig. 4, but for order 6 of SpeX. Note the small residual features in the response spectrum at the locations of Pa12 and Pa13.

and/or air mass difference from any target object. Our experience using SpeX on Mauna Kea indicates that on particularly dry and stable nights, accurate corrections can be achieved when the air mass differences and time intervals between the target and standard observations are as large as 0.1 and 30–60 minutes, respectively. However, for typical conditions, good telluric corrections require closer matches in air mass and shorter time separations.

Accurate flux calibrations require knowledge of the near-infrared fluxes or magnitudes of a set of A0 V stars. Unfortunately, an accurate and comprehensive set of fluxes and/or magnitude measurements for A0 V stars at these wavelengths is currently not available. Therefore, we have chosen to adopt the optical (*B* and *V*) magnitudes given by SIMBAD and the reddening curve given by Rieke & Lebofsky (1985; see also Rieke, Rieke, & Paul 1989) to estimate the flux as a function of wavelength for a given A0 V star. In order to achieve an S/N of about 100 without inordinately long exposure times,

the A0 V star must be reasonably bright; for SpeX observations, A0 V stars with *V* magnitudes of ~ 5 – 8 have been found to work best. SIMBAD provides highly accurate magnitudes for a large number of A0 V stars in this brightness range. In addition, the reddening of such objects is necessarily small, even in the visible. Hence, the precise reddening curve adopted has a minor effect on the results. However, we point out that accurate spectroscopic flux calibrations at near-infrared wavelengths would clearly benefit from a dedicated program of careful *J*, *H*, and *K* observations of a set of A0 V stars. Such a program could also serve as the basis for establishing a set of near-infrared “spectroscopic flux standards” similar to those available in the optical regime.

Proper flux calibration also requires that the seeing and other slit losses do not change substantially between the observations of the target and the A0 V star. If the apertures used for the spectral extraction are wide, moderate changes in seeing should not affect the resulting flux calibrations greatly. In practice, we

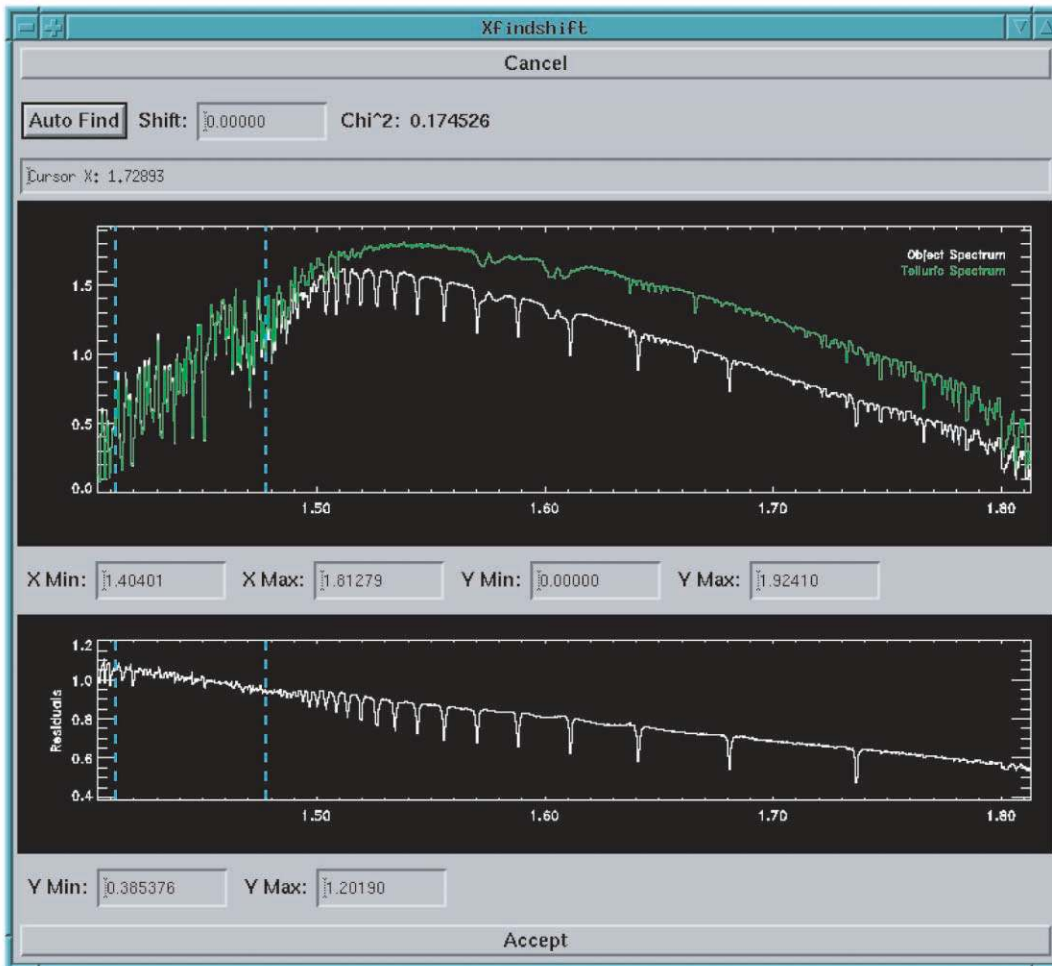


FIG. 6.—Determination of the shift between the telluric correction spectrum and the observed target object spectrum. The vertical dashed lines denote the region used to determine the pixel shift. The solid white line is the observed object spectrum, while the green line is the system response (the inverse of the telluric correction spectrum). The lower panel shows the telluric-corrected spectrum for this order.

find that for typical seeing conditions on Mauna Kea at the IRTF and extraction apertures of $\sim 2''$ width, the derived fluxes agree with those estimated from broadband magnitudes to within a few percent.

4. IMPLEMENTATION: *xtellcor*

The method outlined above (§ 2) has been implemented in a graphical IDL-based code called *xtellcor*. Here we describe the version of *xtellcor* specifically designed to operate on SpeX data. A more general version of *xtellcor* (for non-SpeX data) is similar but does not construct a convolution kernel $K(\lambda)$ from the A0 V star; rather, the user must input values of the parameters describing the IP, $P(\lambda)$. Both versions of the code are available from the SpeX Web site. Help files are included in the distribution packages to guide the user through the various steps.

Starting *xtellcor* brings up an IDL widget containing fields in which the user can enter the information necessary for the

code to run (data directory and filenames for the observed A0 V and target object spectra; see Fig. 2). Since near-infrared magnitudes of a large set of A0 V stars are not available, the user must give the B and V magnitudes of the observed A0 V star. These magnitudes are used to redden and scale the model spectrum of Vega in order to set the flux calibration properly.

At this stage, the user can choose to build a kernel from the observed A0 V star or accept an empirical IP constructed from the profiles of arc lines in SpeX wavelength calibration images. The latter is particularly useful in cases where the IP cannot be constructed from the observed A0 V star (e.g., in the long-wavelength cross-dispersed modes of SpeX that cover the 1.9–5.5 μm range). The arc line profiles have been fitted with an empirical function of the form

$$P(x) = C\{\text{erf}[(x+a)/b] - \text{erf}[(x-a)/b]\}, \quad (13)$$

where C is a normalization constant and a and b determine the

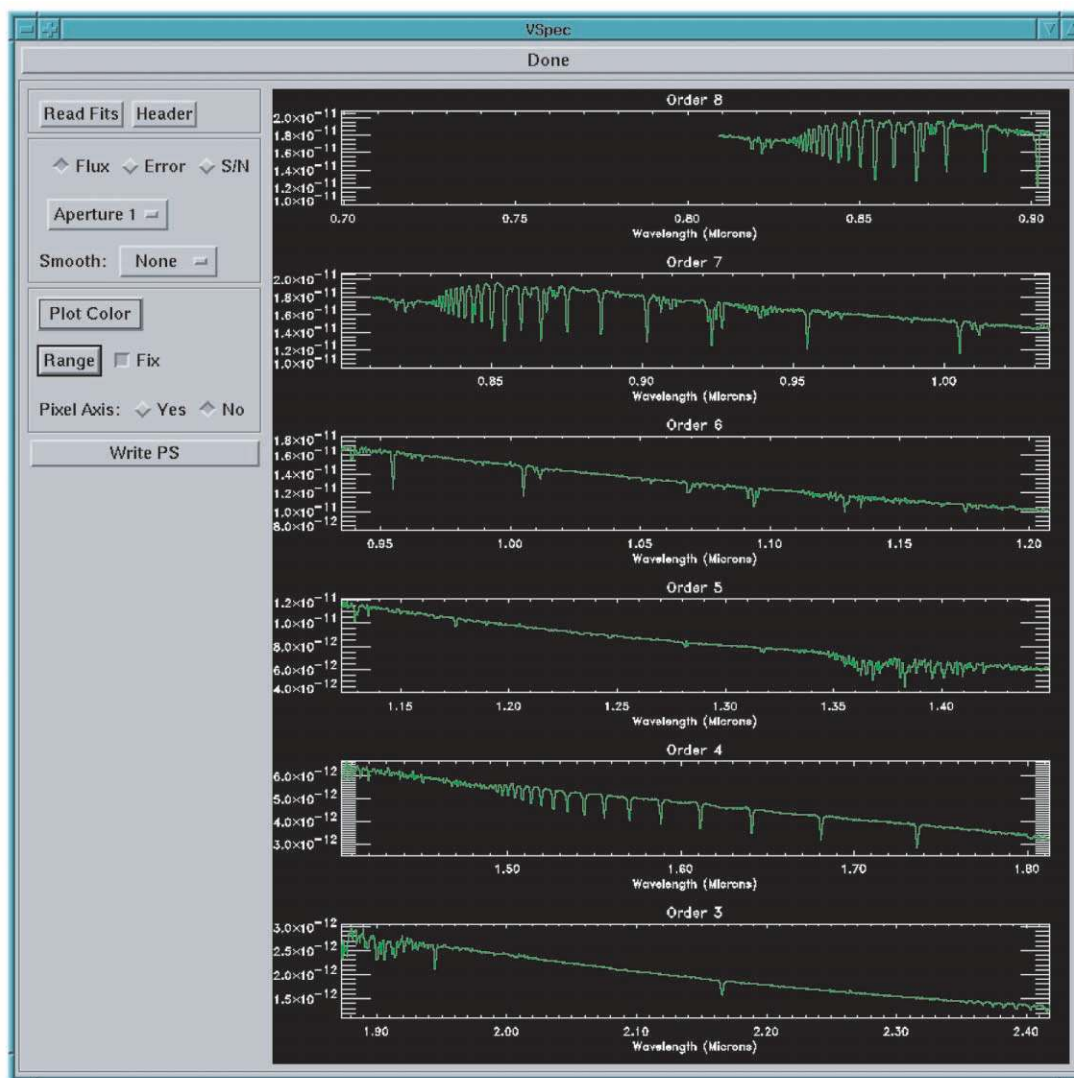


FIG. 7.—Final telluric-corrected spectrum of HD 223385 (A3 Ia).

width of the profile in pixel (x) coordinates. This functional form provides an excellent fit to the observed profiles of the arc lines over all the orders of SpeX and for slit widths of $0''.3$, $0''.5$, $0''.8$, and $1''.6$. By adopting a fit, rather than the observed line profiles themselves, errors in the IP resulting from under-sampling, low S/N, etc., are minimized. The best-fit values of a and b for each slit width are stored in a separate data file.

If the user chooses to build the kernel from the A0 V star, a second window is generated containing the observed spectrum in the requested order. The Pa δ line is located in order 6 of the 0.8–2.4 μm cross-dispersed mode of SpeX. The user can then interactively select continuum regions to be fitted with a polynomial of a desired degree. Once an adequate fit to the continuum in the vicinity of the spectral line is obtained, the normalized spectrum (observed divided by the continuum fit) is displayed (Fig. 3). The user then selects the region of the normalized spectrum to use in the deconvolution, and the code

proceeds to scale both the model line spectrum and the normalized model continuum spectrum by the ratio of the observed equivalent width to the model equivalent width of the selected line (Fig. 3). The relative velocity shift between the observed spectrum and the model is determined and the kernel is generated via deconvolution after the model has been shifted. Our experience has shown that with careful choice of the normalization and deconvolution regions, a kernel can be routinely constructed that yields maximum deviations of the convolved model from the observed line profile of less than 2% over the selected wavelength range. We typically achieve maximum deviations of less than 1.5% and rms deviations of less than 0.75% across the Pa δ line for data with $\lambda/\Delta\lambda = 2000$ and $S/N \geq 100$.

The observed spectrum of the A0 V star is divided by the full scaled and shifted model after the latter has been convolved with the kernel (either the semiempirical IP or one built from the observed A0 V star) and the result is displayed for each

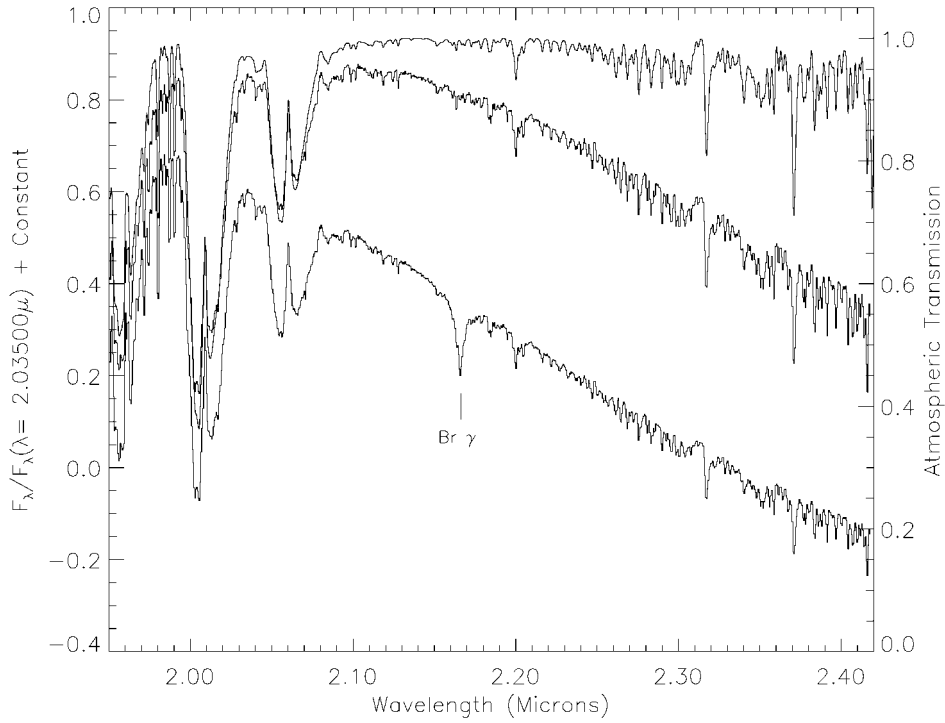


Fig. 8a

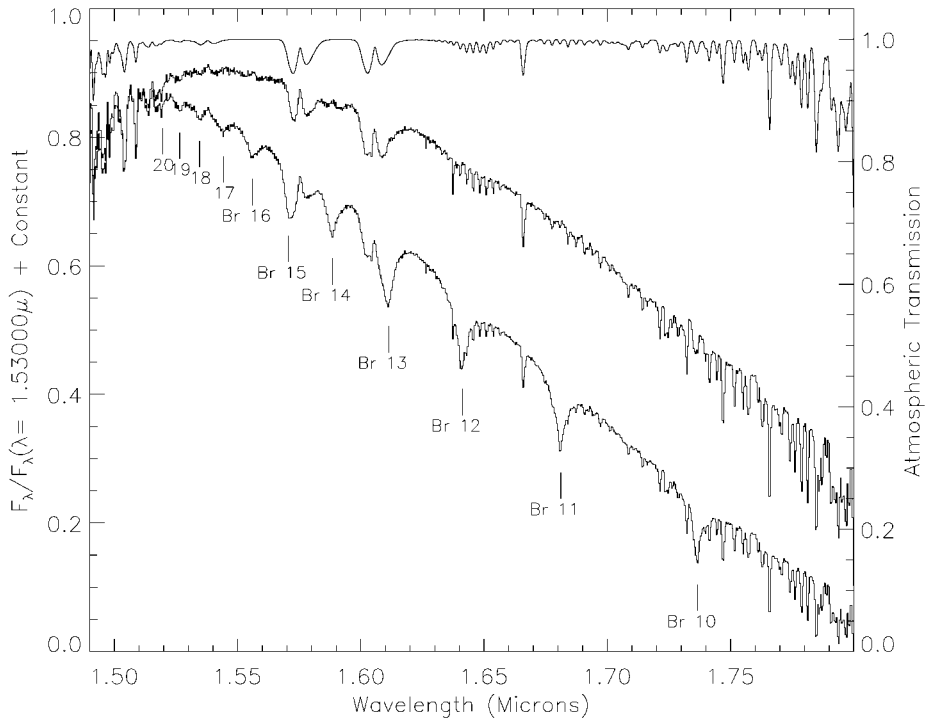


Fig. 8b

FIG. 8.—Response curves generated by *xtellcor* for the various short-wavelength cross-dispersed orders of SpeX (*a*: order 3; *b*: order 4; *c*: order 5; *d*: order 6; *e*: order 7; *f*: order 8). In each plot, the curves shown (*bottom to top*) are the observed spectrum of an A0 V star (HD 223386, $V = 6.328$), the response curve generated with *xtellcor*, and a theoretical atmospheric transmission spectrum computed for typical conditions on Mauna Kea, an air mass similar to that of the observations, and a resolving power of 2000. The intrinsic stellar H lines of the Paschen and Brackett series are identified. Small residuals are seen in the response curves at the locations of the Pa11–Pa15 lines. Note that the theoretical atmosphere curve does not extend below $0.85 \mu\text{m}$.

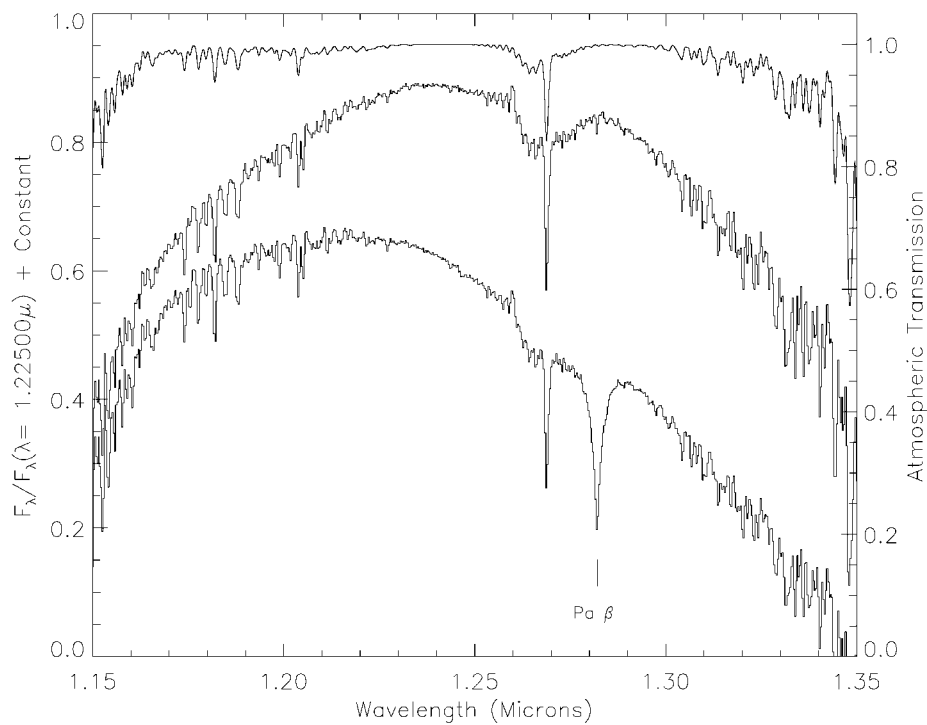


Fig. 8c

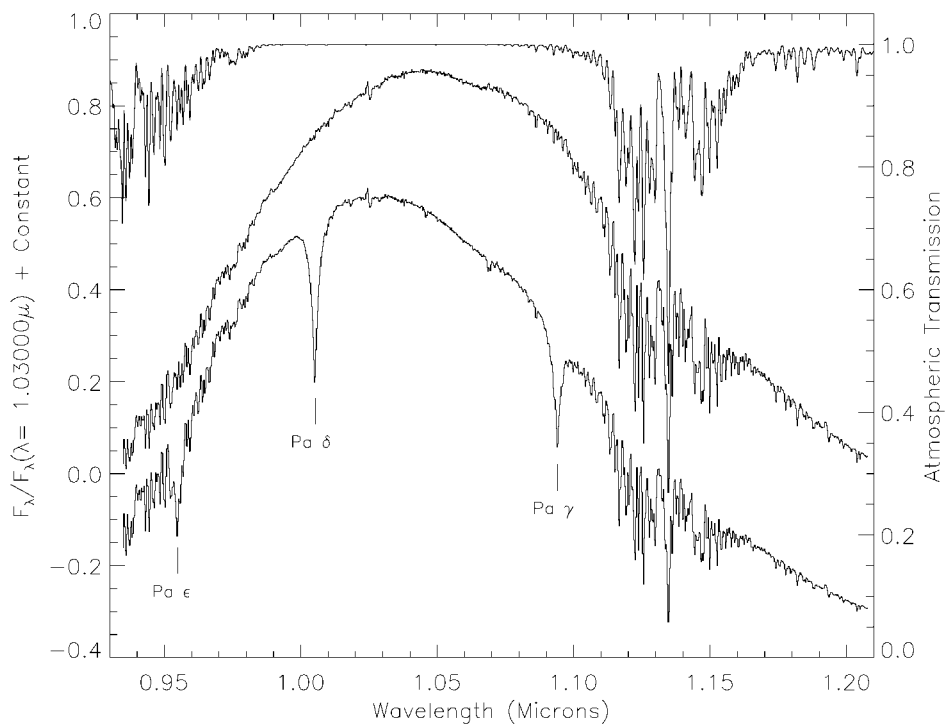


Fig. 8d

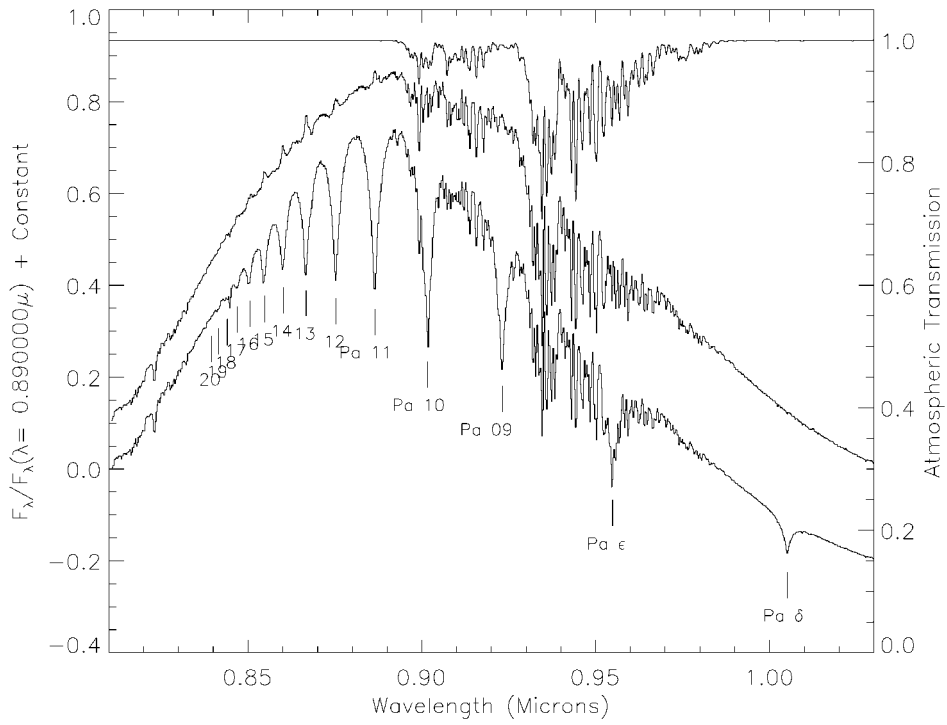


Fig. 8e

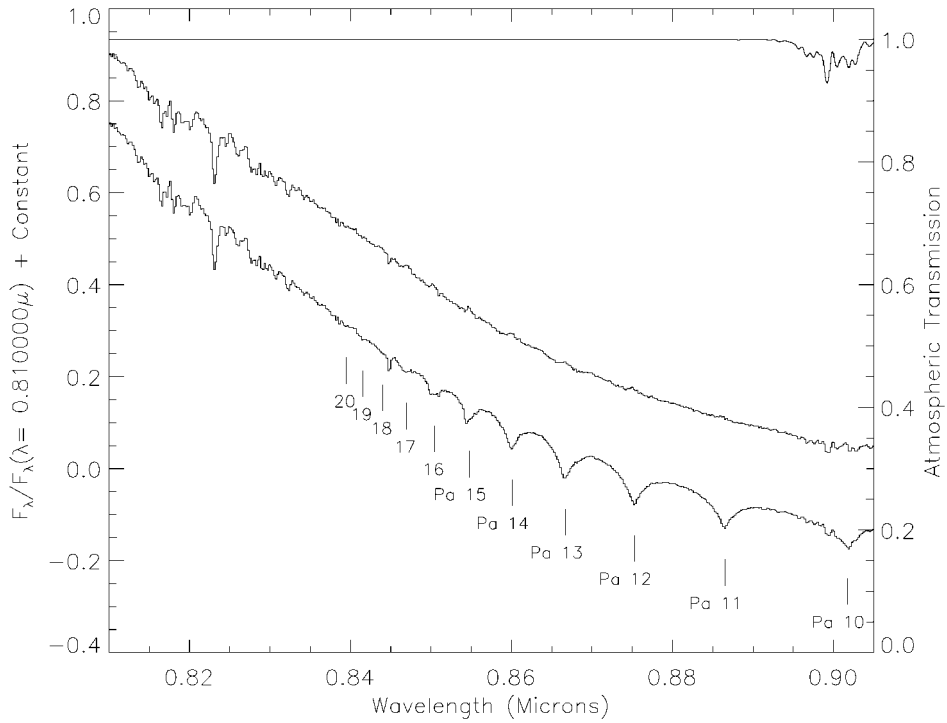


Fig. 8f

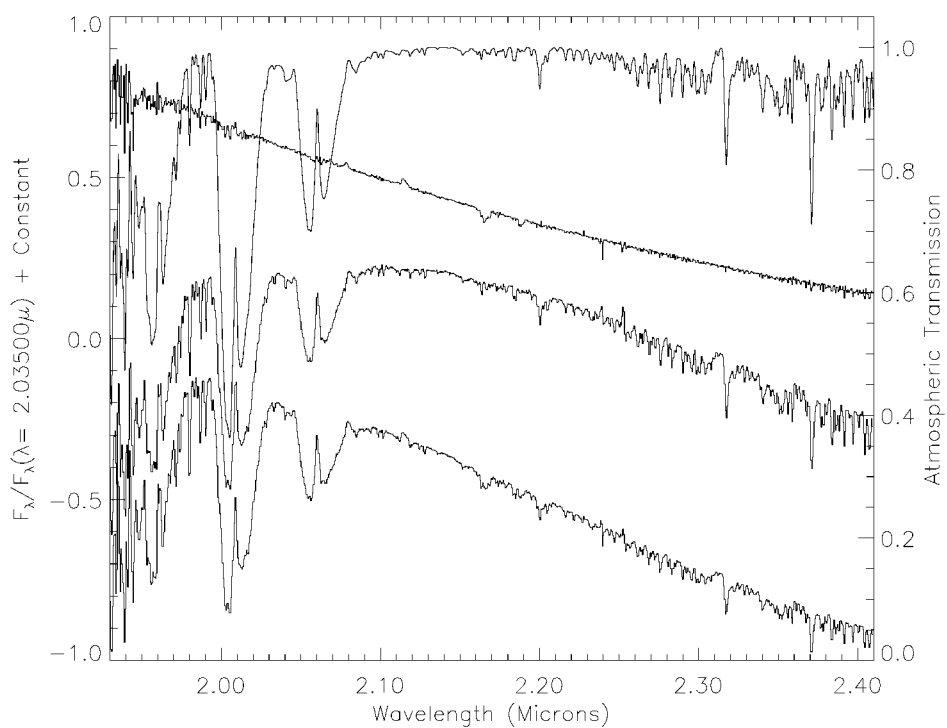


Fig. 9a

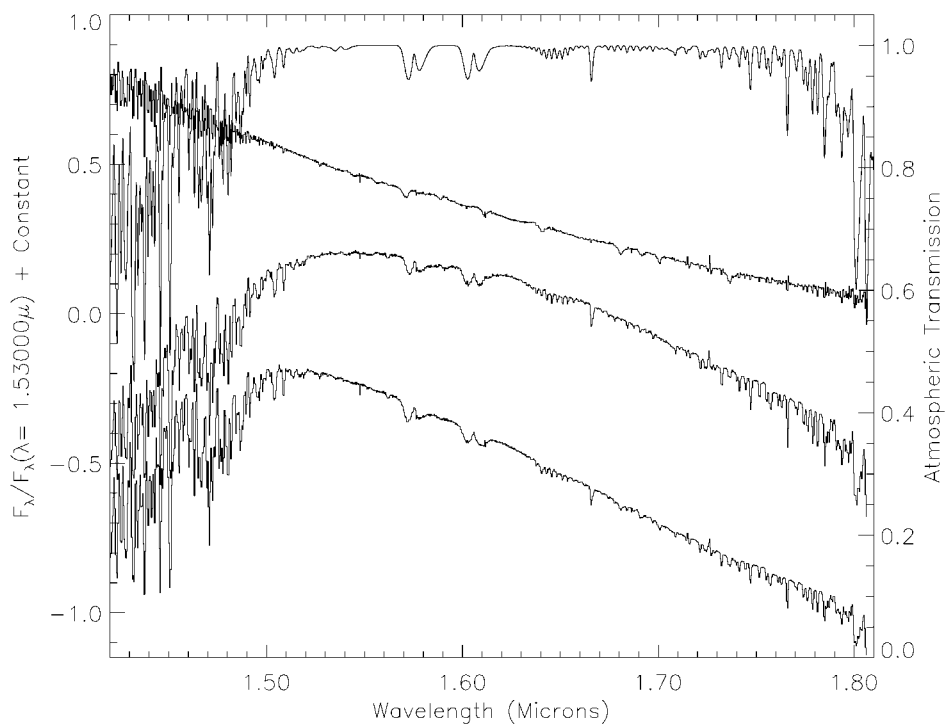


Fig. 9b

FIG. 9.—Final telluric-corrected spectra of an O6.5 V((f)) star (HD 206267), generated by *xtellcor* for the various short-wavelength cross-dispersed orders of SpeX (*a*: order 3; *b*: order 4; *c*: order 5; *d*: order 6; *e*: order 7). In each plot, the spectra are ordered from bottom to top as follows: the observed stellar O6.5 V((f)) spectrum, the response curve constructed by *xtellcor* from the A0 V telluric standard, the telluric-corrected O6.5 V((f)) spectrum, and the representative theoretical atmospheric transmission spectrum (for reference, right-hand axis).

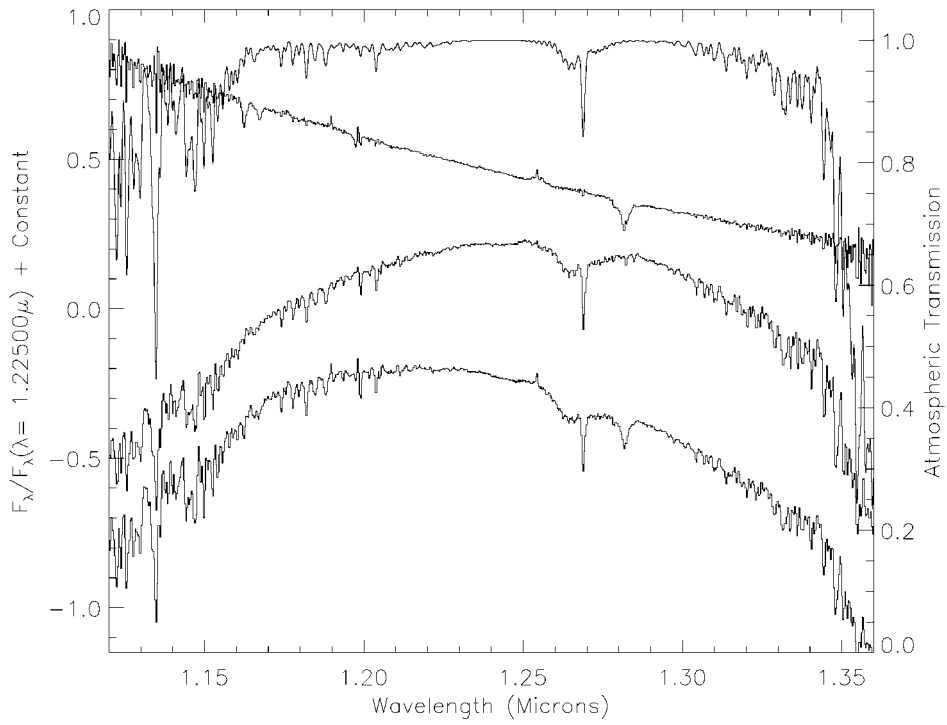


Fig. 9c

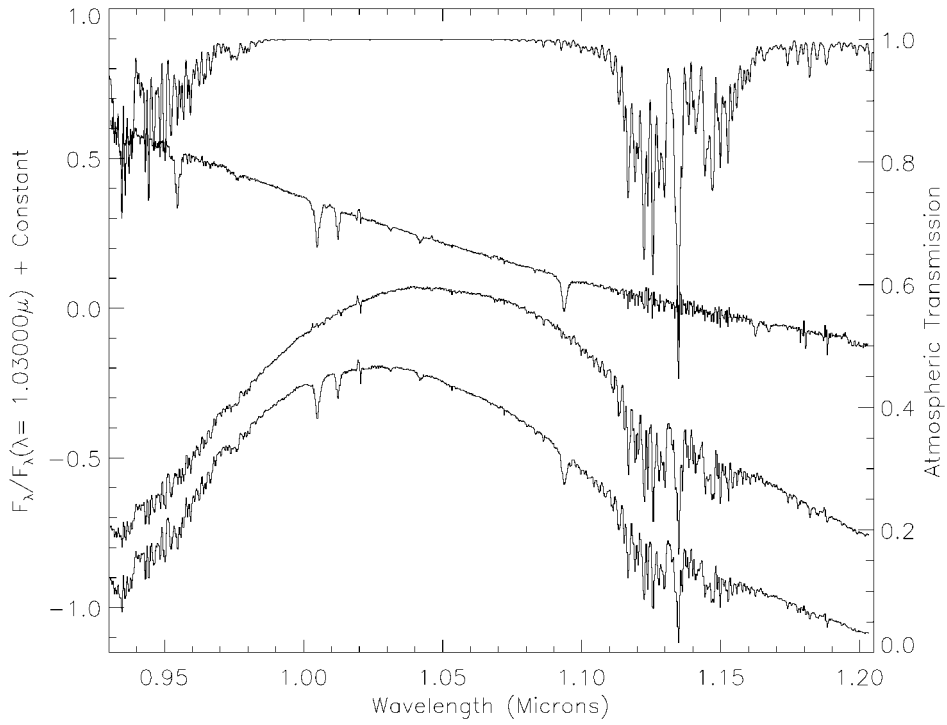


Fig. 9d

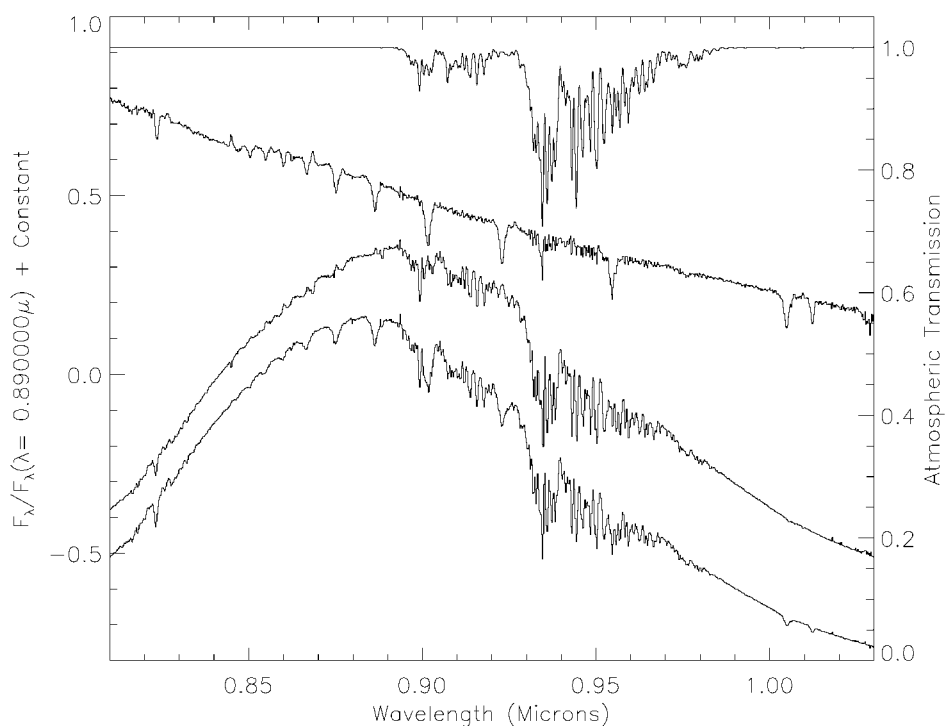


Fig. 9e

order (Figs. 4 and 5). This is the first version of the response spectrum. At this stage, the user can adjust the velocity shift of the model and modify the strengths of the various H lines to better match those of the observed A0 V star. Mismatches in line strengths and velocities are easily seen in the telluric correction spectrum as “absorption” or “emission” features at the location of the H lines. The user can adjust the individual H line strengths of the model by varying the scale factors applied to the model line equivalent widths until residual features are minimized in the response spectrum in the regions around the lines (see Figs. 4 and 5). This can be done graphically using the mouse. A representative spectrum of the atmospheric transmission is displayed to guide the user in this task. (As an additional aid to the user, this atmospheric transmission curve can be multiplied by the estimated throughput curve of SpeX.) The code also includes a procedure to determine the scale factor for a single isolated line automatically. This procedure requires two points on either side of a line, fits a Gaussian plus a continuum to the response spectrum, and adjusts the strength of the model line until the height of the Gaussian is minimized. Most of the H lines can be easily removed from the response/telluric correction spectrum at this stage. However, we have found that the Pa12, Pa13, and Pa14 lines (and occasionally the Pa11 and Pa15 lines) almost always leave residual features in the telluric correction spectrum that cannot be eliminated, no matter how the line strengths are scaled. This

indicates that the Vega model we are using does not provide a good match to the profiles of these lines in the spectra of real A0 V stars. Once an improved Vega model is available, we will incorporate it into the code. Fortunately, even in the vicinity of these discrepant lines, the residual features are generally less than a few percent and therefore do not result in large errors in the final spectrum.

The resulting set of line-strength scale factors are joined with a spline fit to generate a line scaling spectrum that is multiplied by the line model. The normalized continuum (divided by the spline fit) is also multiplied by this line scaling curve. This has the effect of modifying the jumps seen in the continuum model. Since the highest order lines rarely require modification of their strengths, the jumps are normally scaled by the equivalent-width ratio determined from the line used to build the kernel. The full, scaled and shifted, model is then used to regenerate the response spectrum. When the user is satisfied with the modifications to the model, the final telluric correction spectrum is generated. This can then be shifted in pixels in order to align the telluric absorption features with those seen in the target object spectrum. This procedure should minimize any wavelength shifts due to flexure between the observations of the target object and the standard star. The code automatically computes the best pixel shift using a region of the spectra selected by the user (Fig. 6). The shifts are usually less than 0.5 pixel for standard stars observed within 10° of the target

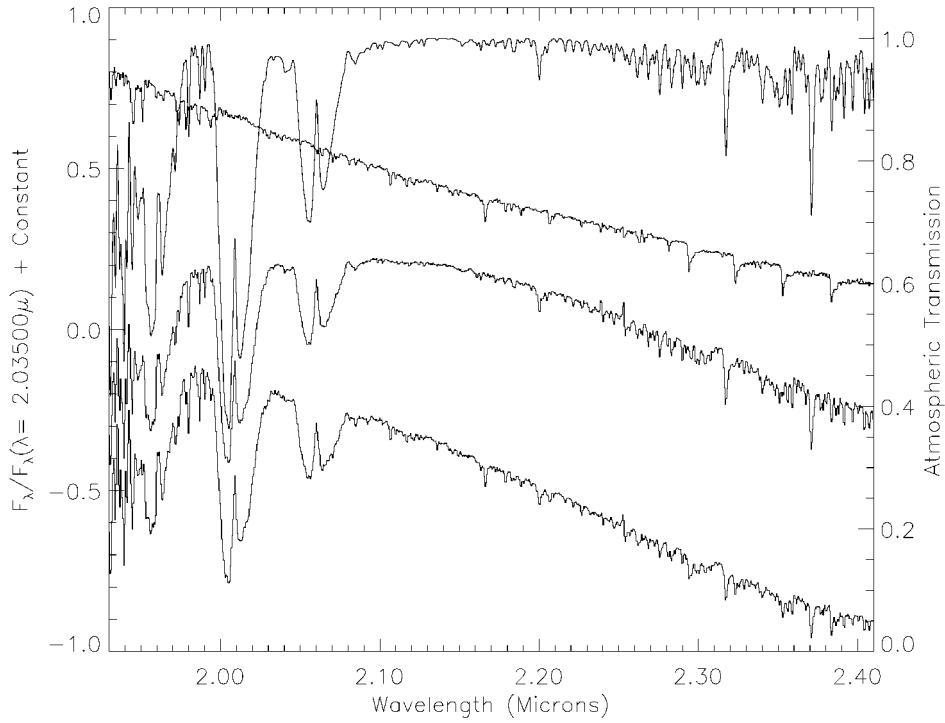


Fig. 10a

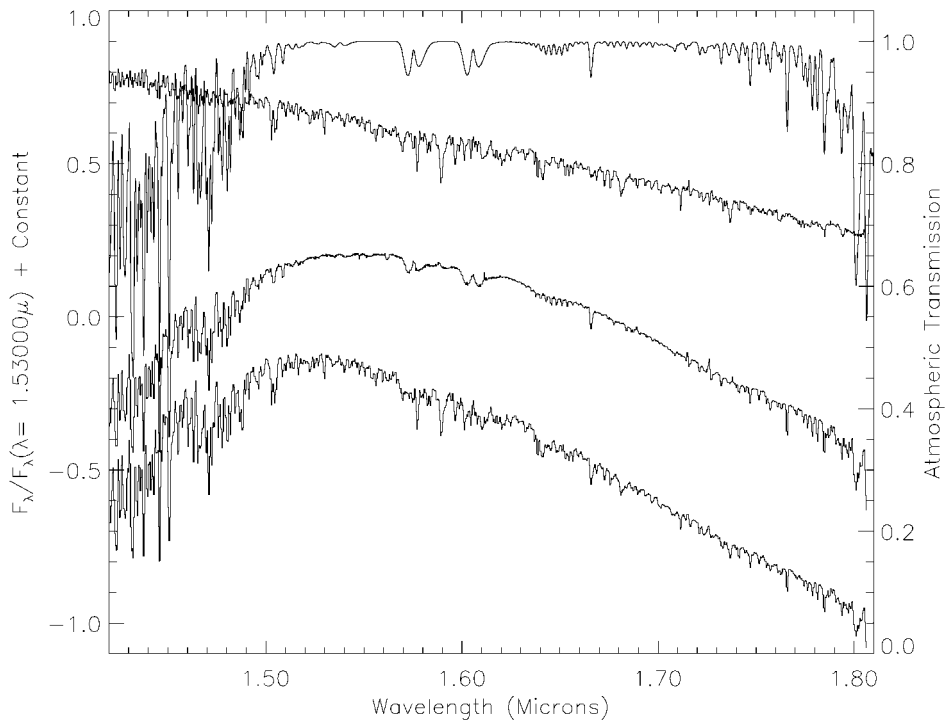


Fig. 10b

FIG. 10.—Final telluric-corrected spectra of a G8 IIIa star (HD 16139), generated by *xtellcor* for the various short-wavelength cross-dispersed orders of SpeX (*a*: order 3; *b*: order 4; *c*: order 5; *d*: order 6; *e*: order 7). In each plot, the spectra are ordered from bottom to top as follows: the observed stellar G8 IIIa spectrum, the response curve constructed by *xtellcor* from the A0 V telluric standard, the telluric-corrected G8 IIIa spectrum, and the representative theoretical atmospheric transmission spectrum (for reference, right-hand axis).

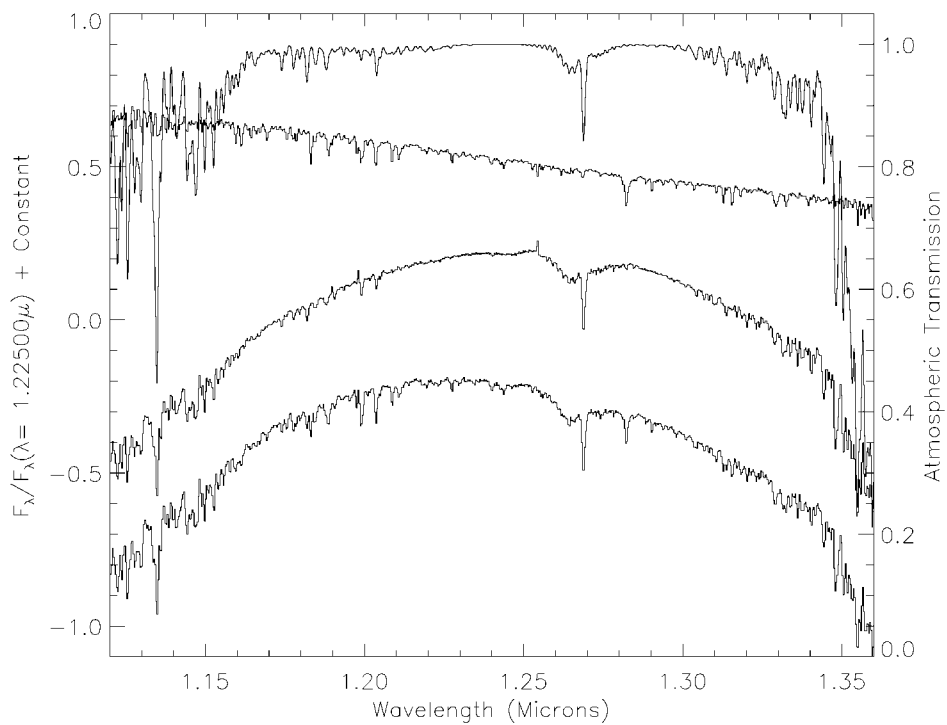


Fig. 10c

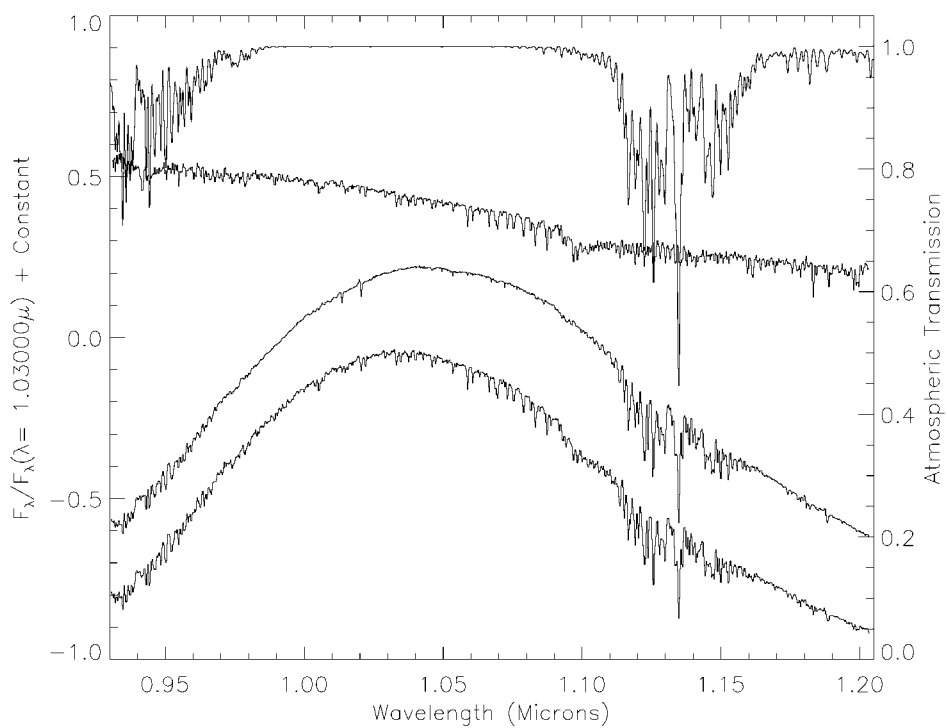


Fig. 10d

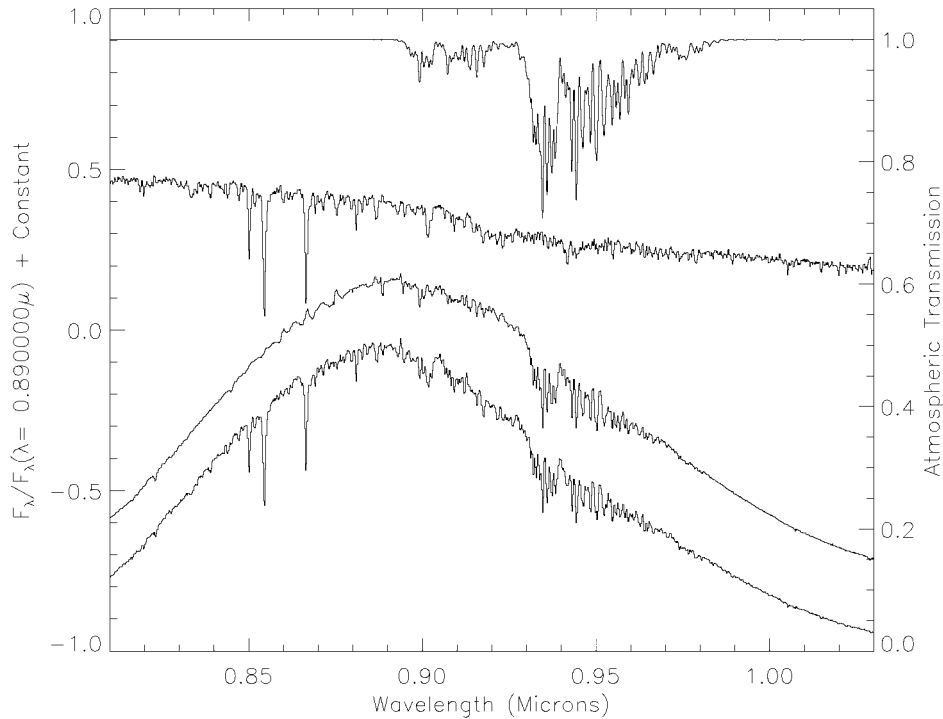


Fig. 10e

object. The shift value is then applied to all the orders of the telluric correction spectrum. Finally, the observed target object spectrum is multiplied by the shifted telluric correction spectrum, order by order, and the final corrected object spectrum is generated (Fig. 7). The user has a choice of units for the final output spectrum. The user can also choose to save both the telluric correction spectrum itself and the modified Vega spectrum. The former is useful if a single A0 V star has been observed as a telluric standard for multiple target objects; a separate routine allows this previously constructed telluric correction spectrum to be read in and used to correct the observed spectra of other objects.

5. APPLICATION AND EXAMPLES

The response spectra generated by *xtellcor* for the various short-wavelength cross-dispersed orders of SpeX are shown in Figures 8a–8f and Figure 1 of Cushing et al. (2003). As can be seen from these figures, the features seen in the response spectra match extremely well with those seen in the theoretical atmospheric spectrum, and the stellar H lines seen in the observed A0 V spectra are removed to a very high accuracy. Small residuals can be seen in Figure 8e at the locations of the Pa11–Pa15 lines. Aside from these features, however, the accuracy of the telluric corrections is limited only by the S/N of

the input A0 V spectrum. The latter is determined by the exposure time, flat-fielding errors, instrument throughput, and the atmospheric transmission itself. (Clearly, the derived telluric correction spectrum will be highly uncertain in those wavelength regions where the atmosphere is so opaque that the observed spectra have S/N close to zero.) Of course, as explained above, the accuracy of the final corrected target spectra also depends on the differences in time and air mass between the observations of the A0 V star and the target object, as well as the S/N of the target spectrum.

Examples of typical output spectra generated by *xtellcor* can be seen in Figure 7 (for HD 223385, A3 Ia), Figure 9 [for HD 206267, O6.5 V((f))], and Figure 10 (for HD 16139, G8 IIIa), as well as in the papers by Cushing et al. (2003) and Rayner et al. (2003). The differences in air mass and time of observations between the objects shown in these figures and their telluric A0 V standards were ≤ 0.05 and ≤ 20 minutes, respectively, in all cases. However, the atmospheric conditions (humidity, seeing, transparency) were variable during the observations. The S/N values of the A0 V spectra were always ≥ 100 . As the figures demonstrate, the method enables us to recover the target spectrum even in regions of relatively low atmospheric transmission. We note that the equivalent widths of the intrinsic stellar absorption and emission features seen in the spectra in Figures 7 and 9 agree well with those measured for these spectral classes by Hanson et al. (1996) in the *K* band.

6. SUMMARY

We have described a technique for correcting observed near-infrared medium-resolution spectra for telluric absorption. Designed specifically for data obtained with SpeX at the IRTF, the method makes use of a high S/N (≥ 100) spectrum of an A0 V star observed near in time and close in air mass to the object of interest and uses a theoretical model of Vega to remove the intrinsic H absorption features seen in the stellar spectrum. The method can regularly generate telluric correction spectra accurate to better than 2% in the vicinity of the intrinsic stellar H lines and significantly better than that at other wavelengths.

We have demonstrated the application of this method on spectra obtained with SpeX. A set of graphical, easy-to-use, IDL-based routines (part of a general SpeX spectral reduction package called Spextool_extension version 1.1), implements the procedures described in this paper and can be found on the

SpeX Web page at the IRTF Web site.⁵ A general version of the code, useful for performing telluric corrections on spectra obtained with other near-infrared spectrographs (with $\lambda/\Delta\lambda < 50,000$), is also available. In addition, a form that generates a list of the A0 V stars listed in the SIMBAD database near a set of specified coordinates is also posted on the Web site.

We thank R. Kurucz for providing the high-resolution model of Vega and J. Tonry for suggestions regarding the form of the instrument profiles. We also thank the SpeX team for their expert construction of an invaluable instrument and the staff at the IRTF for their assistance at the telescope. This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France.

⁵ <http://irtfweb.ifa.hawaii.edu/Facility/spex>.

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