

A PLATEAU IN THE PLANET POPULATION BELOW TWICE THE SIZE OF EARTH

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ABSTRACT

We carry out an independent search of *Kepler* photometry for small transiting planets with sizes 0.5–8.0 times that of Earth and orbital periods between 5 and 50 days, with the goal of measuring the fraction of stars harboring such planets. We use a new transit search algorithm, TERRA, optimized to detect small planets around photometrically quiet stars. We restrict our stellar sample to include the 12,000 stars having the lowest photometric noise in the *Kepler* survey, thereby maximizing the detectability of Earth-size planets. We report 129 planet candidates having radii less than $6 R_E$ found in 3 years of *Kepler* photometry (quarters 1–12). Forty-seven of these candidates are not in Batalha et al. (2012), which only analyzed photometry from quarters 1–6. We gather Keck HIRES spectra for the majority of these targets leading to precise stellar radii and hence precise planet radii. We make a detailed measurement of the completeness of our planet search. We inject synthetic dimmings from mock transiting planets into the actual *Kepler* photometry. We then analyze that injected photometry with our TERRA pipeline to assess our detection completeness for planets of different sizes and orbital periods. We compute the occurrence of planets as a function of planet radius and period, correcting for the detection completeness as well as the geometric probability of transit, R_*/a . The resulting distribution of planet sizes exhibits a power law rise in occurrence from $5.7 R_E$ down to $2 R_E$, as found in Howard et al. (2012). That rise clearly ends at $2 R_E$. The occurrence of planets is consistent with constant from $2 R_E$ toward $1 R_E$. This unexpected plateau in planet occurrence at $2 R_E$ suggests distinct planet formation processes for planets above and below $2 R_E$. We find that $15.1^{+1.8}_{-2.7}\%$ of solar type stars—roughly one in six—has a 1– $2 R_E$ planet with $P = 5$ –50 days.

Subject headings: techniques: photometric — planetary systems

1. INTRODUCTION

The *Kepler* Mission has discovered an extraordinary sample of more than 2300 planets with radii ranging from larger than Jupiter to smaller than Earth (Borucki et al. 2011b; Batalha et al. 2012). Cleanly measuring and debiasing this distribution will be one of *Kepler*'s great legacies. Howard et al. (2012), H12 hereafter, took a key step, showing that the planet radius distribution increases substantially with decreasing planet size down to at least $2 R_E$. While the distribution of planets of all periods and radii contains a wealth of information, we choose to focus on the smallest planets. Currently, only *Kepler* is able to make quantitative statements about the occurrence of planets down to $1 R_E$.

The occurrence distributions in H12 were based on planet candidates detected in the first four months of *Kepler* photometry (Borucki et al. 2011b). These planet candidates were detected by a sophisticated pipeline developed by the *Kepler* team Science Operations Center (Twicken et al. 2010; Jenkins et al. 2010a).¹ Under-

standing pipeline completeness, the fraction of planets missed by the pipeline as a function of size and period, is a key component to measuring planet occurrence. Pipeline completeness can be assessed by injecting mock dimmings into photometry and measuring the rate at which injected signals are found. The completeness of the official *Kepler* pipeline has yet to be measured in this manner. This was the key reason why H12 were cautious interpreting planet occurrence under $2 R_E$.

In this work, we focus on determining the occurrence of small planets. To maximize our sensitivity to small planets, we restrict our stellar sample to include only the 12,000 stars having the lowest photometric noise in the *Kepler* survey. We comb through quarters 1–12 (Q1–Q12) — 3 years of *Kepler* photometry — with a new algorithm, TERRA, optimized to detect low signal-to-noise transit events. We determine TERRA's sensitivity to planets of different periods and radii by injecting synthetic transits into *Kepler* photometry and measuring the recovery rate as a function of planet period and radius.

We describe our selection of 12,000 low-noise targets

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¹ Since H12, Batalha et al. (2012) added many candidates, bringing the number of public KOIs to > 2300 . In addition,

the *Kepler* team planet search pipeline has continued to evolve (Smith et al. 2012; Stumpe et al. 2012).

in Section 2. We comb their photometry for exoplanet transits with **TERRA**, introduced in Section 3. We report candidates found with **TERRA** (Section 4), which we combine with our measurement of pipeline completeness (Section 5) to produce debiased measurements of planet occurrence (Section 6). We offer some comparisons between **TERRA** planet candidates and those from Batalha et al. (2012) in Section 7 as well as occurrence measured using both catalogs in Section 8. We offer some interpretations of the constant occurrence rate for planets smaller than $2 R_E$ in Section 9.

2. THE BEST12K STELLAR SAMPLE

We restrict our study to the best 12,000 solar type stars from the perspective of detecting transits by Earth-size planets, hereafter, the “Best12k” sample. For the smallest planets, uncertainty in the occurrence distribution stems largely from pipeline incompleteness due to the low signal-to-noise ratio (SNR) of an Earth-size transit.

Our initial sample begins with the 102,835 stars that were observed during every quarter from Q1–Q9.² From this sample, following H12, we select 73,757 “solar subset” stars that are solar-type G and K having $T_{\text{eff}} = 4100\text{--}6100$ K and $\log g = 4.0\text{--}4.9$ (cgs). T_{eff} and $\log g$ values are present in the *Kepler* Input Catalog (KIC; Brown et al. 2011) which is available online.³ Figure 1 shows the KIC-based T_{eff} and $\log g$ values as well as the solar subset. KIC stellar parameters have large uncertainties: $\sigma(\log g) \sim 0.4$ dex and $\sigma(T_{\text{eff}}) \sim 200$ K (Brown et al. 2011). As we will discuss in Section 4, we determine stellar parameters for the majority of **TERRA** planet candidates spectroscopically. For the remaining cases, we use stellar parameters that were determined photometrically, but incorporated a main sequence prior (Batalha et al. 2012). After refining the stellar parameters, we find that 10 of the 129 **TERRA** planet candidates fall outside of the $T_{\text{eff}} = 4100\text{--}6100$ K and $\log g = 4.0\text{--}4.9$ (cgs) solar subset.

From the 73,757 stars that pass our cuts on $\log g$ and T_{eff} , we choose the 12,000 lowest noise stars. *Kepler* target stars have a wide range of noise properties, and there are several ways of quantifying photometric noise. The *Kepler* team computes quantities called CDPP3, CDPP6, and CDPP12, which are measures of the photometric scatter in 3, 6, and 12 hour bins (Jenkins et al. 2010a). Since CDPP varies by quarter, we adopt the maximum 6-hour CDPP over Q1–Q9 as our nominal noise metric. We use the maximum noise level (as opposed to median or mean) because a single quarter of noisy photometry can set a high noise floor for planet detection. One may circumvent this problem by removing noisy regions of photometry, which is a planned upgrade to **TERRA**. Figure 2 shows the distribution of $\max(\text{CDPP6})$ among the 73,757 stars considered for our sample.

In choosing our sample, we wanted to include stars amenable to the detection of planets as small as $1 R_E$. We picked the 12,000 quietest stars based on preliminary completeness estimates. The noisiest star in the Best12k sample has $\max(\text{CDPP6})$ of 79.2 ppm. We es-

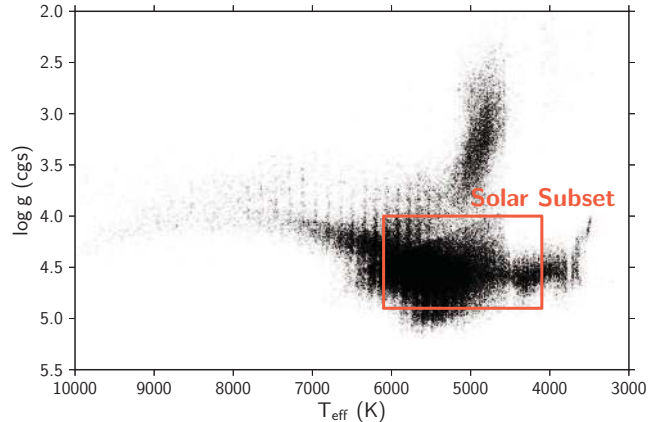


Figure 1. *Kepler* target stars observed every quarter from Q1–Q9. The rectangle marks the “solar subset” of stars with $T_{\text{eff}} = 4100\text{--}6100$ K and $\log g = 4.0\text{--}4.9$ (cgs).

timated that the ~ 100 ppm transit of an Earth-size planet would be detected at $\text{SNR}_{\text{CDPP}} \sim 1.25$.⁴ Given that Q1–Q12 contains roughly 1000 days of photometry, we expected to detect a 5-day planet at $\text{SNR}_{\text{CDPP}} \sim 1.25 \times \sqrt{1000/5} \sim 18$ (a strong detection) and to detect a 50-day planet at $\text{SNR}_{\text{CDPP}} \sim 1.25 \times \sqrt{1000/50} \sim 5.6$ (a marginal detection). In our detailed study of completeness, described in Section 5, we find that **TERRA** recovers most planets down to $1 R_E$ having $P = 5\text{--}50$ days.

We draw stars from the H12 solar subset for two reasons. First, we may compare our planet occurrence to that of H12 without the complication of varying occurrence with different stellar types. We recognize that subtle differences may exist between the H12 and Best12k stellar sample. One such difference is that the Best12k is noise-limited, while the H12 sample is magnitude-limited. H12 included bright stars with high photometric variability, which are presumably young and/or active stars. Planet formation efficiency could depend on stellar age. Planets may be less common around older stars that formed before the metallicity of the Galaxy was enriched to current levels. This work assesses planet occurrence for a set of stars that are systematically selected to be 3–10 Gyr old by virtue of their reduced magnetic activity.

The second reason for adopting the H12 solar subset is a practical consideration of our completeness study. As shown in Section 5, we parameterize pipeline efficiency as a function of P and R_P . Because M-dwarfs have smaller radii than G-dwarfs, an Earth-size planet dims an M-dwarf more substantially and should be easier for **TERRA** to detect. Thus, measuring completeness as a function of P , R_P , and R_* (or perhaps P and R_P/R_*) is appropriate when analyzing stars of significantly different sizes. Such extensions are beyond the scope of this paper, and we consider stars with $R_* \sim R_\odot$.

3. PLANET SEARCH PIPELINE

Identifying the smallest transiting planets in *Kepler* photometry requires a sophisticated automated pipeline. Our pipeline is called “**TERRA**” and consists of three major

² We ran **TERRA** on Q1–Q12 photometry, but we selected the Best12k sample before Q10–Q12 were available.

³ <http://archive.stsci.edu/Kepler/kic.html>

⁴ SNR_{CDPP} , the expected SNR using the $\max(\text{CDPP6})$ metric, is different from the SNR introduced in Section 3.2. SNR_{CDPP} is more similar to the SNR computed by the *Kepler* team, which adopts $\text{SNR}_{\text{CDPP}} > 7.1$ as their detection threshold.

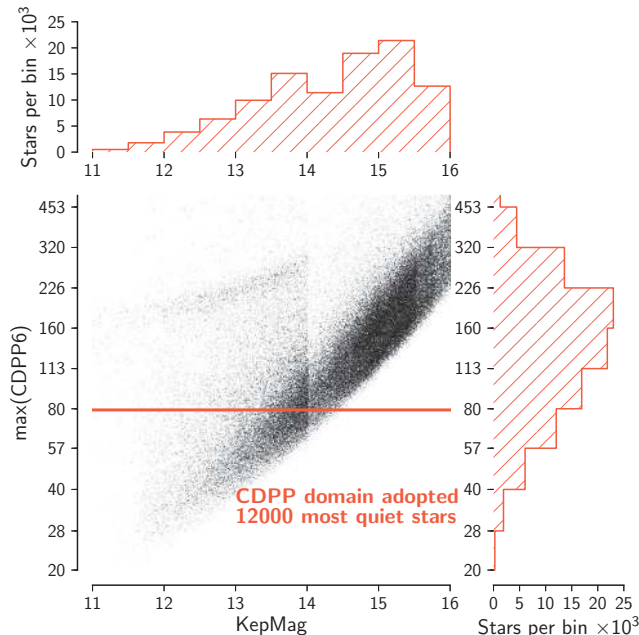


Figure 2. Stellar photometric noise level plotted against *Kepler* magnitude. Noise level is the maximum value of CDPP6 over Q1–Q9. Of the 73,757 stars that pass our cuts on T_{eff} and $\log g$, we select the 12,000 most quiet stars. The line shows $\text{max}(\text{CDPP6}) = 79.2$ ppm, corresponding to the noisiest star in the Best12k sample, well below the median value of 143 ppm.

components. First, TERRA calibrates photometry in the time domain. Then, TERRA combs the calibrated photometry for periodic, box-shaped signals by evaluating the signal-to-noise ratio (SNR) over a finely-spaced grid in transit period (P), epoch (t_0) and duration (ΔT). Finally, TERRA fits promising signals with a Mandel & Agol (2002) transit model and rejects signals that are not consistent with an exoplanet transit. We review the calibration component in Section 3.1, but refer the reader to Petigura & Marcy (2012) for a detailed description. We present, for the first time, the grid-search and light curve fitting components in Sections 3.2 and 3.3.

3.1. Photometric Calibration

We briefly review the major time domain components of TERRA; for a more complete description, please refer to Petigura & Marcy (2012). We begin with *Kepler* “simple aperture long cadence photometry,” which we downloaded from the Mikulski Archive for Space Telescopes (MAST). This photometry is the total photoelectrons accumulated within a predefined target aperture over a 29.4 minute interval (Fraquelli & Thompson 2012). We remove thermal settling events manually and cosmic rays using a median filter. Next, we remove photometric trends longer than 10 days with a high-pass filter. Finally, we identify photometric modes shared by a large ensemble of stars with using a robust principal components analysis. The optimum linear combination of the four most significant modes is removed from each light curve individually.

3.2. Grid-based Transit Search

We then search for periodic, box-shaped signals in ensemble-calibrated photometry. Such a search involves evaluating the SNR over a finely sampled grid in period

(P), epoch (t_0), and duration (ΔT), i.e.

$$\text{SNR} = \text{SNR}(P, t_0, \Delta T). \quad (1)$$

Our approach is similar to the widely-used BLS algorithm of Kovács et al. (2002) as well as to the TPS component of the *Kepler* pipeline (Jenkins et al. 2010b). BLS, TPS, and TERRA are all variants of a “matched filter” (North 1943). The way in which such an algorithm searches through P , t_0 , and ΔT is up to the programmer. We choose to search first through ΔT (outer loop), then P , and finally, t_0 (inner loop).

For computational simplicity, we consider transit durations that are integer numbers of long cadence measurements. Since we search for transits with $P = 5\text{--}50$ days, we try $\Delta T = [3, 5, 7, 10, 14, 18]$ long cadence measurements, which span the range of expected transit durations, 1.5 to 8.8 hours, for G and K dwarf stars.

After choosing ΔT , we compute the mean depth, $\overline{\delta F}(t_i)$, of a putative transit with duration = ΔT centered at t_i for each cadence. $\overline{\delta F}$ is computed via

$$\overline{\delta F}(t_i) = \sum_j F(t_{i-j}) G_j \quad (2)$$

where $F(t_i)$ is the median-normalized stellar flux at time t_i and G_j is the j th element of the following kernel

$$\mathbf{G} = \frac{1}{\Delta T} \left[\underbrace{\frac{1}{2}, \dots, \frac{1}{2}}_{\Delta T}, \underbrace{-1, \dots, -1}_{\Delta T}, \underbrace{\frac{1}{2}, \dots, \frac{1}{2}}_{\Delta T} \right]. \quad (3)$$

As an example, if $\Delta T = 3$,

$$\mathbf{G} = \frac{1}{3} \left[\frac{1}{2}, \frac{1}{2}, \frac{1}{2}, -1, -1, -1, \frac{1}{2}, \frac{1}{2}, \frac{1}{2} \right]. \quad (4)$$

We search over a finely sampled grid of trial periods from 5–50 days and epochs ranging from t_{start} to $t_{\text{start}} + P$, where t_{start} is the time of the first photometric observation. For a given $(P, t_0, \Delta T)$ there are N_T putative transits with depths $\overline{\delta F}_i$, for $i = 0, 1, \dots, N_T - 1$. For each $(P, t_0, \Delta T)$ triple, we compute SNR from

$$\text{SNR} = \frac{\sqrt{N_T}}{\sigma} \text{mean}(\overline{\delta F}_i), \quad (5)$$

where σ is a robust estimate (median absolute deviation) of the noise in bins of length ΔT .

For computational efficiency, we employ the “Fast Folding Algorithm” (FFA) of Staelin (1969) as implemented in Petigura & Marcy (2013; in prep.). Let $P_{\text{cad},0}$ be a trial period that is an integer number of long cadence measurements, e.g. $P_{\text{cad},0} = 1000$ implies $P = 1000 \times 29.4 \text{ min} = 20.43$ days. Let $N_{\text{cad}} = 51413$ be the length of the Q1–Q12 time series measured in long cadences. Leveraging the FFA, we compute SNR at the following periods:

$$P_{\text{cad},i} = P_{\text{cad},0} + \frac{i}{M-1}; \quad i = 0, 1, \dots, M-1 \quad (6)$$

where $M = N_{\text{cad}}/P_{\text{cad},0}$ rounded up to the nearest power of two. In our search from 5–50 days, $P_{\text{cad},0}$ ranges from 245–2445, and we evaluate SNR at $\sim 10^5$ different periods. At each $P_{\text{cad},i}$ we evaluate SNR for $P_{\text{cad},0}$ different

starting epochs. All told, for each star, we evaluate SNR at $\sim 10^9$ different combinations of P , t_0 , and ΔT .

Due to runtime and memory constraints, we store only one SNR value for each of the trial periods. TERRA stores the maximum SNR at that period for all ΔT and t_0 . We refer to this one-dimensional distribution of SNR as the ‘‘SNR periodogram,’’ and we show the KIC-3120904 SNR periodogram in Figure 3 as an example. Because we search over many ΔT and t_0 at each trial period, fluctuations often give rise SNR ~ 8 events and set the detectability floor in the SNR periodogram. For KIC-3120904, a star not listed in the Batalha et al. (2012) planet catalog, we see a SNR peak of 16.6, which rises clearly above stochastic background.

If the maximum SNR in the SNR periodogram exceeds 12, we pass that particular $(P, t_0, \Delta T)$ on to the ‘‘data validation’’ (DV) step, described in the following section, for additional vetting. We chose 12 as our SNR threshold by trial and error. Note that the median absolute deviation of many samples drawn from a Gaussian distribution is 0.67 times the standard deviation, i.e. $\sigma_{\text{MAD}} = 0.67\sigma_{\text{STD}}$. Therefore, TERRA SNR = 12 corresponds roughly to SNR = 8 in a BLS or TPS search.

Since TERRA only passes the $(P, t_0, \Delta T)$ triple with the highest SNR on to DV, TERRA does not detect additional planets with lower SNR due to either smaller size or longer orbital period. As an example of TERRA’s insensitivity to small candidates in multi-candidate systems, we show the TERRA SNR periodogram for KIC-5094751 in Figure 4. Batalha et al. (2012) lists two candidates belonging to KIC-5094751: KOI-123.01 and KOI-123.02 with $P = 6.48$ and 21.22 days, respectively. Although the SNR periodogram shows two sets of peaks coming from two distinct candidates, TERRA only identifies the first peak. Automated identification of multi-candidate systems is a planned upgrade for TERRA. Another caveat is that TERRA assumes strict periodicity and struggles to detect low SNR transits with significant transit timing variations, i.e. variations longer than the transit duration.

3.3. Data Validation

If the SNR periodogram has a maximum SNR peak > 12 , we flag the corresponding $(P, t_0, \Delta T)$ for additional vetting. Following the language of the official *Kepler* pipeline, we refer to these triples as ‘‘threshold crossing events’’ (TCEs), since they have high photometric SNR, but are not necessarily consistent with an exoplanet transit. TERRA vets the TCEs in a step called ‘‘data validation,’’ again following the nomenclature of the official *Kepler* pipeline. Data validation (DV), as implemented in the official *Kepler* pipeline, is described in Jenkins et al. (2010a). We emphasize that TERRA DV does not depend on the DV component of the *Kepler* team pipeline.

We show the distribution of maximum SNR for each Best12k star in Figure 5. Among the Best12k stars, 738 have a maximum SNR peak exceeding 12. Adopting SNR = 12 as our threshold balances two competing needs: the desire to recover small planets (low SNR) and the desire to remove as many non-transit events as possible before DV (high SNR). As discussed below, only 129 out of all 738 events with SNR > 12 are consistent with an exoplanet transit, with noise being responsible for the remaining 609. As shown in Figure 5, that number grows

rapidly as we lower the SNR threshold. For example, the number of TCEs grows to 3055 with a SNR threshold of 10, dramatically increasing the burden on the DV component.

A substantial number (347) of TCEs are due to harmonics or subharmonics of TCEs outside of the $P = 5\text{--}50$ day range and are discarded. In order to pass DV, a TCE must also pass a suite of four diagnostic metrics. The metrics are designed to test whether a light curve is consistent with an exoplanet transit. We describe the four metrics in Table 1 along with the criteria the TCE must satisfy in order pass DV. The metrics and cuts were determined by trial and error. We recognize that the TERRA DV metrics and cuts are not optimal and discard a small number of compelling exoplanet candidates, as discussed in Section 7.3. However, since we measure TERRA’s completeness by injection and recovery of synthetic transits, the sub-optimal nature of our metrics and cuts is incorporated into our completeness corrections.

Our suite of automated cuts removes all but 145 TCEs. We perform a final round of manual vetting and remove 16 additional TCEs, leaving 129 planet candidates. Most TCEs that we remove manually come from stars with highly non-stationary photometric noise properties. Some stars have small regions of photometry that exceed typical noise levels by a factor of 3. We show the SNR periodogram for one such star, KIC-7592977, in Figure 6. Our definition of SNR (Equation 5) incorporates a single measure of photometric scatter based on the median absolute deviation, which is insensitive to short bursts of high photometric variability. In such stars, fluctuations readily produce SNR ~ 12 events and raise the detectability floor to SNR ~ 12 , up from SNR ~ 8 in most stars. We also visually inspect phase-folded light curves for coherent out-of-transit variability, not caught by our automated cuts, and for evidence of a secondary eclipse.

4. SMALL PLANETS FOUND BY TERRA

Out of the 12,000 stars in the Best12k sample, TERRA detected 129 planet candidates achieving SNR > 12 that passed our suite of DV cuts as well as visual inspection. Table 2 lists the 129 planet candidates. We derive planet radii using R_P/R_\star (from Mandel-Agol model fits) and R_\star from spectroscopy (when available) or broadband photometry.

We obtained spectra for 100 of the 129 stars using HIRES (Vogt et al. 1994) at the Keck I telescope with standard configuration of the California Planet Survey (Marcy et al. 2008). These spectra have resolution of $\sim 50,000$, at a signal-to-noise of 45 per pixel at 5500 Å. We determine stellar parameters using a routine called *SpecMatch* (Howard et al. 2013, in prep). In brief, *SpecMatch* compares a stellar spectrum to a library of ~ 800 spectra with $T_{\text{eff}} = 3500\text{--}7500$ K and $\log g = 2.0\text{--}5.0$ (determined from LTE spectral modeling). Once the target spectrum and library spectrum are placed on the same wavelength scale, we compute χ^2 , the sum of the squares of the pixel-by-pixel differences in normalized intensity. The weighted mean of the ten spectra with the lowest χ^2 values is taken as the final value for the effective temperature, stellar surface gravity, and metallicity. We estimate *SpecMatch*-derived stellar radii are uncertain to 10% RMS, based on tests of stars having known

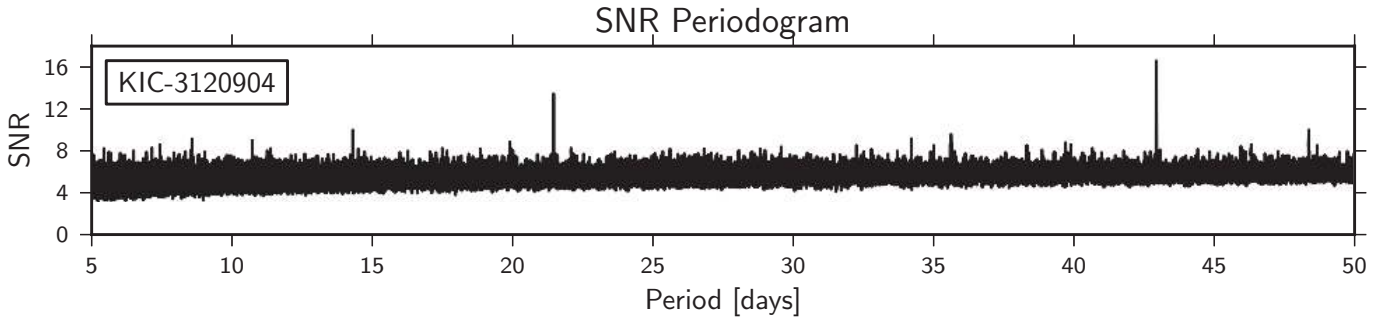


Figure 3. SNR periodogram of KIC-3120904 photometry. We evaluate SNR over a finely-spaced, three-dimensional grid of P , t_0 , and ΔT . We store the maximum SNR for each trial period, resulting in a one-dimensional distribution of SNR. A planet candidate (not in Batalha et al. 2012) produces a SNR peak of 16.6 at $P = 42.9$ days, which rises clearly above the detection floor of $\text{SNR} \sim 8$.

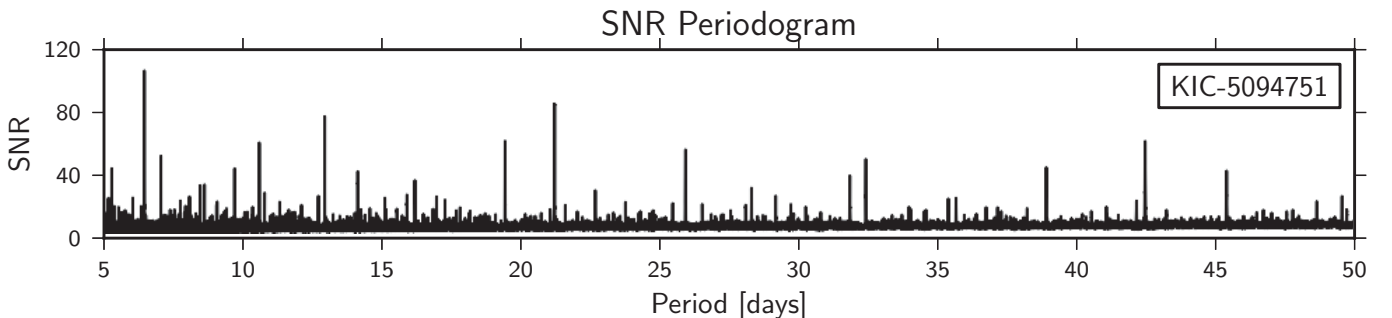


Figure 4. SNR periodogram of KIC-5094751 photometry, demonstrating **TERRA**’s insensitivity to lower SNR candidates in multi-candidate systems. Batalha et al. (2012) lists two planets belonging to KIC-5094751, KOI-123.01 and KOI-123.02 with $P = 6.48$ and 21.22 days, respectively. **TERRA** detected KOI-123.01 with a period of 6.48 days (highest SNR peak). Sub-harmonics belonging to KOI-123.01 are visible at $[2, 3, \dots] \times P = [13.0, 19.4, \dots]$ days. A second set of SNR peaks due to KOI-123.02 ($P = 21.2$ days) is visible at $[0.5, 2, \dots] \times P = [10.6, 42.4, \dots]$ days. Had we removed the transit due to KOI-123.01, KOI-123.02 would be easily detectable due its high SNR of ~ 80 . **TERRA** does not yet include multi-candidate logic and is thus blind to lower SNR candidates in multi-candidate systems.

Table 1

In order to pass the “data validation” (DV) stage, a “threshold crossing event” (TCE) must pass the following suite of cuts.

name	description	value
<code>s2n_out_on_in</code>	Compelling transits have flat out-of-transit light curves. For a TCE with $(P, t_0, \Delta T)$, we remove the transit region from the light curve and evaluate the SNR of all other $(P', t'_0, \Delta T')$ triples where $P = P'$ and $\Delta T = \Delta T'$. <code>s2n_out_on_in</code> is the ratio of the two highest SNR events.	< 0.7
<code>med_on_mean</code>	Since our definition of SNR (Equation 5) depends on the arithmetic mean of individual transit depths, outliers occasionally produce high SNR TCEs. For each TCE, we compute a robust SNR,	> 1.0
$\text{medSNR} = \frac{\sqrt{N_T}}{\sigma} \text{median}(\delta F_i).$		
<code>autor</code>	<code>med_on_mean</code> is medSNR divided by SNR as defined in Equation 5. We compute the circular autocorrelation of the phase-folded light curve. <code>autor</code> is the ratio of the highest autocorrelation peak (at 0 lag) to the second highest peak and is sensitive to out-of-transit variability.	> 1.6
<code>taur</code>	We fit the phase-folded light curve with a Mandel & Agol (2002) model. <code>taur</code> is the ratio of the best fit transit duration to the maximum duration given the KIC stellar parameters and assuming a circular orbit.	< 2.0

radii from high resolution spectroscopy and asteroseismology.

For 27 stars where spectra are not available, we adopt the photometrically-derived stellar parameters of Batalha et al. (2012). These parameters are taken from the KIC (Brown et al. 2011), but then modified so that they lie on the Yonsei-Yale stellar evolution models of Demarque et al. (2004). The resulting stellar radii have uncertainties of 35% (rms), but can be incor-

rect by a factor of 2 or more. As an extreme example, the interpretations of the three planets in the KOI-961 system (Muirhead et al. 2012) changed dramatically when HRES spectra showed the star to be an M5 dwarf ($0.2 R_\odot$ as opposed to $0.6 R_\odot$ listed in the KIC). We could not obtain spectra for two stars, KIC-7345248 and KIC-8429668, which were not present in Batalha et al. (2012). We determine stellar parameters for these stars by fitting the KIC photometry to Yonsei-

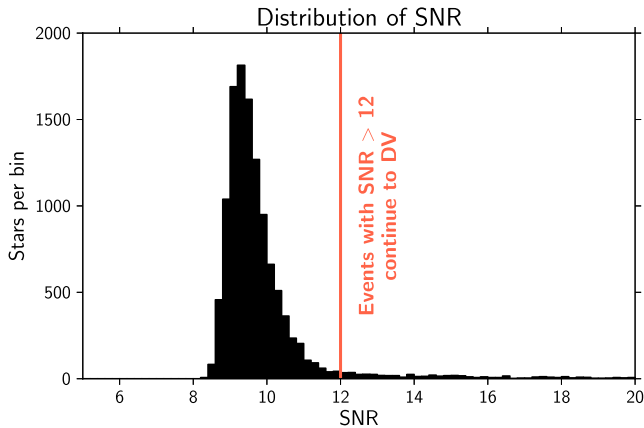


Figure 5. Distribution of the highest SNR peak for each star in the Best12k sample. We show SNR = 5–20 to highlight the distribution of low SNR events. The 738 stars with SNR > 12 are labeled “threshold crossing events” (TCEs) and are subjected to additional scrutiny in the “data validation” component of TERRA.

Yale stellar models. We adopt 35% fractional errors on photometrically-derived stellar radii.

Once we determine P and t_0 , we fit a Mandel & Agol (2002) model to the phase-folded photometry. Such a model has three free parameters: R_P/R_* , the planet to stellar radius ratio; τ , the time for the planet to travel a distance R_* during transit; and b , the impact parameter. In this work, R_P/R_* is the parameter of interest. However, b and R_P/R_* are covariant, i.e. a transit with b approaching unity only traverses the limb of the star, and thus produces a shallower transit depth. In order to account for this covariance, best fit parameters were computed via Markov Chain Monte Carlo. We find that the fractional uncertainty on R_P/R_* , $\frac{\sigma(R_P/R_*)}{R_P/R_*}$ can be as high as 10%, but is generally less than 5%. Therefore, the error on R_P due to covariance with b is secondary to the uncertainty on R_* .

We show the distribution of TERRA candidates in Figure 7 over the two-dimensional domain of planet radius and orbital period. Our 129 candidates range in size from $6.83 R_E$ to $0.48 R_E$ (smaller than Mars). The median TERRA candidate size is $1.58 R_E$. In Figure 8, we show the substantial overlap between the TERRA planet sample and those produced by the *Kepler* team. TERRA recovers 82 candidates listed in Batalha et al. (2012). We discuss the significant overlap between the two works in detail in Section 7. As of August 8th, 2012, 10 of our TERRA candidates were listed as false positives in an internal database of *Kepler* planet candidates maintained by Jason Rowe (Jason Rowe, 2012, private communication) and are shown as blue crosses in Figure 8. We do not include these 10 candidates in our subsequent calculation of occurrence. Table 2 lists the KIC identifier, best fit transit parameters, stellar parameters, planet radius, and *Kepler* team false positive designation of all 129 candidates revealed by the TERRA algorithm. The best fit transit parameters include orbital period, P ; time of transit center, t_0 ; planet to star radius ratio, R_P/R_* ; time for planet to cross R_* during transit, τ ; and impact parameter, b . We list the following stellar properties: effective temperature, T_{eff} ; surface gravity, $\log g$; and stellar ra-

dius, R_* .

5. COMPLETENESS OF PLANET CATALOG

When measuring the distribution of planets as a function of P and R_P , understanding the number of missed planets is as important as finding planets themselves. H12 accounted for completeness in a rough sense based on signal-to-noise considerations. For each star in their sample, they estimated the SNR over a range of P and R_P using CDPP as an estimate of the photometric noise on transit-length timescales. H12 chose to accept only planets with SNR > 10 in a single quarter of photometry for stars brighter than $K_p = 15$. This metric used CDPP and was a reasonable pass on the data, particularly when the pipeline completeness was unknown. Determining expected SNR from CDPP does not incorporate the real noise characteristics of the photometry, but instead approximates noise on transit timescales as stationary (CDDP assumed to be constant over a quarter) and Gaussian distributed. Moreover, identifying small transiting planets with transit depths comparable to the noise requires a complex, multistage pipeline. Even if the integrated SNR is above some nominal threshold, the possibility of missed planets remains a concern.

We characterize the completeness of our pipeline by performing an extensive suite of injection and recovery experiments. We inject mock transits into raw photometry, run this photometry through the same pipeline used to detect planets, and measure the recovery rate. This simple, albeit brute force, technique captures the idiosyncrasies of the TERRA pipeline that are missed by simple signal-to-noise considerations.

We perform 10,000 injection and recovery experiments using the following steps:

1. We select a star randomly from the Best12k sample.
2. We draw (P, R_P) randomly from log-uniform distributions over 5–50 days and 0.5 – $16.0 R_E$.
3. We draw impact parameter and orbital phase randomly from uniform distributions ranging from 0 to 1.
4. We generate a Mandel & Agol (2002) model.
5. We inject it into the “simple aperture photometry” of the selected star.

We then run the calibration, grid-based search, and data validation components of TERRA (Sections 3.1, 3.2, and 3.3) on this photometry and calculate the planet recovery rate. We do not, however, perform the visual inspection described in Section 3.3. An injected transit is considered recovered if the following two criteria are met: (1) The highest SNR peak passes all DV cuts and (2) the output period and epoch are consistent with the injected period and epoch to within 0.01 and 0.1 days, respectively.

Figure 9 shows the distribution of recovered simulations as a function of period and radius. Nearly all simulated planets with $R_P > 1.4 R_E$ are recovered, compared to almost none with $R_P < 0.7 R_E$. Pipeline completeness is determined in small bins in (P, R_P) -space by dividing the number of successfully recovered transits by the total

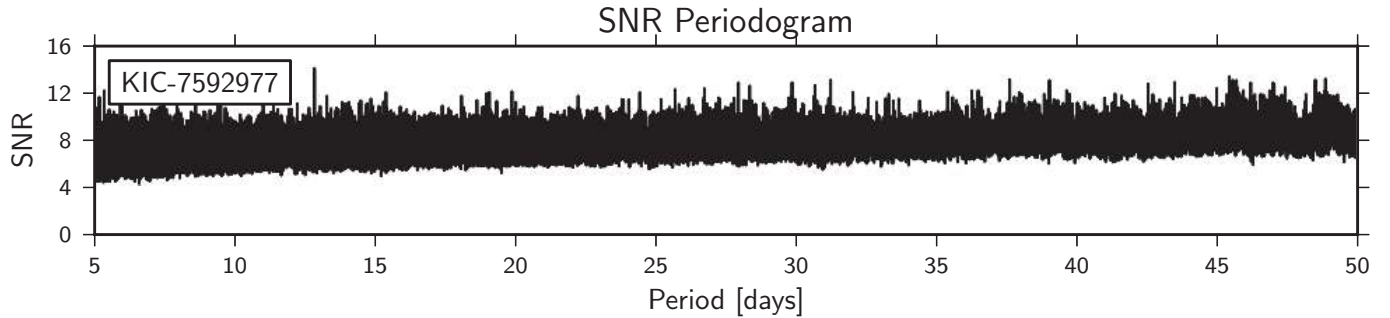


Figure 6. SNR periodogram of KIC-7592977, which passed the automated DV cuts, but was removed manually. KIC-7592977 photometry exhibited short bursts of high photometric scatter, which raised the noise floor to SNR ~ 12 , up from SNR ~ 8 as in most stars.

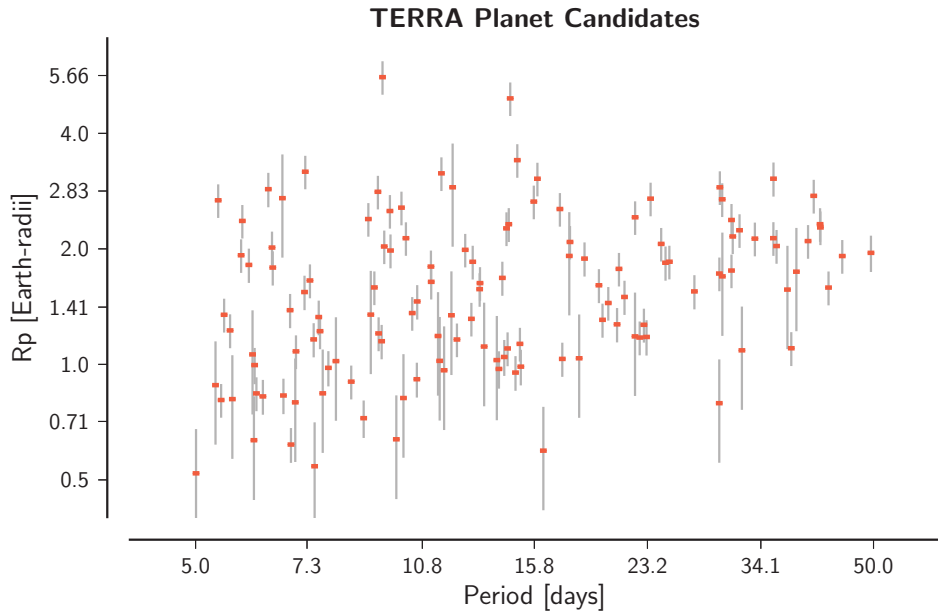


Figure 7. Periods and radii of 129 planet candidates detected by TERRA. Errors on R_P are computed via $\frac{\sigma(R_P)}{R_P} = \sqrt{\left(\frac{\sigma(R_*)}{R_*}\right)^2 + \left(\frac{\sigma(R_P/R_*)}{R_P/R_*}\right)^2}$, where R_P/R_* is the radius ratio. The error in R_P stems largely from the uncertainty in stellar radii. We adopt $\frac{\sigma(R_*)}{R_*} = 10\%$ for the 100 stars with spectroscopically determined R_* and $\frac{\sigma(R_*)}{R_*} = 35\%$ for the remaining stars with R_* determined from photometry. Using MCMC, we find the uncertainty in R_P/R_* is generally $< 5\%$ and thus a minor component of the overall error budget.

number of injected transits in a bin-by-bin basis. This ratio is TERRA’s recovery rate of putative planets within the Best12k sample. Thus, our quoted completeness estimates only pertain to the low photometric noise Best12k sample. Had we selected an even more rarified sample, e.g. the “Best6k,” the region of high completeness would extend down toward smaller planets.

6. OCCURRENCE OF SMALL PLANETS

Following H12, we define planet occurrence, f , as the fraction of a defined population of stars having planets within a domain of planet radius and period, including all orbital inclinations. TERRA, however, is only sensitive to one candidate (highest SNR) per system, so we report occurrence as the fraction of stars with *one or more* planets with $P = 5\text{--}50$ days. Our occurrence measurements apply to the Best12k sample of low-noise, solar-type stars described in Section 2.

In computing planet occurrence in the Best12k sample, we follow the prescription in H12 with minor modifications. Notably, we have accurate measures of detection completeness described in the previous section. In contrast, H12 estimated completeness based on the presumed signal-to-noise of the transit signal, suffering both from approximate characterization of photometric noise using CDPP and from poor knowledge of the efficiency of the planet-finding algorithm for all periods and sizes.

For each P - R_P bin, we count the number of planet candidates, $n_{\text{pl,cell}}$. Each planet that transits represents many that do not transit given the orientation of their orbital planes with respect to *Kepler*’s line of sight. Assuming random orbital alignment, each observed planet represents a/R_* total planets when non-transiting geometries are considered. For each cell, we compute the number of augmented planets, $n_{\text{pl,aug,cell}} = \sum_i a_i/R_{*,i}$, which accounts for planets with non-transiting geome-

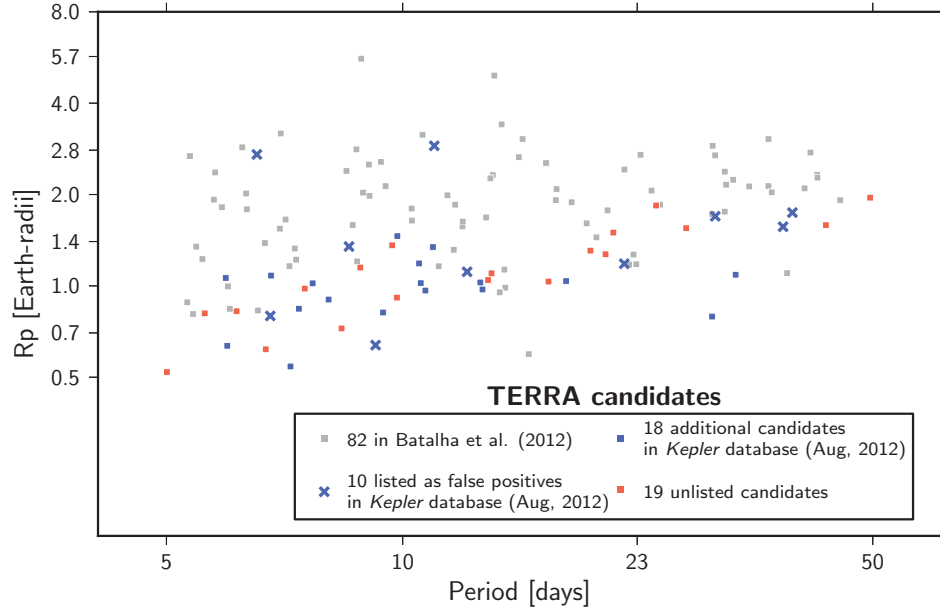


Figure 8. Periods and radii of all 129 TERRA planet candidates. The gray points show candidates that were listed in Batalha et al. (2012). The blue crosses represent candidates deemed false positives by the *Kepler* team as of August 8, 2012 (Jason Rowe, private communication 2012). These false positives are removed from our sample prior to computing occurrence. Eighteen additional candidates were listed in the same *Kepler* team database. Red points show 19 unlisted TERRA candidates.

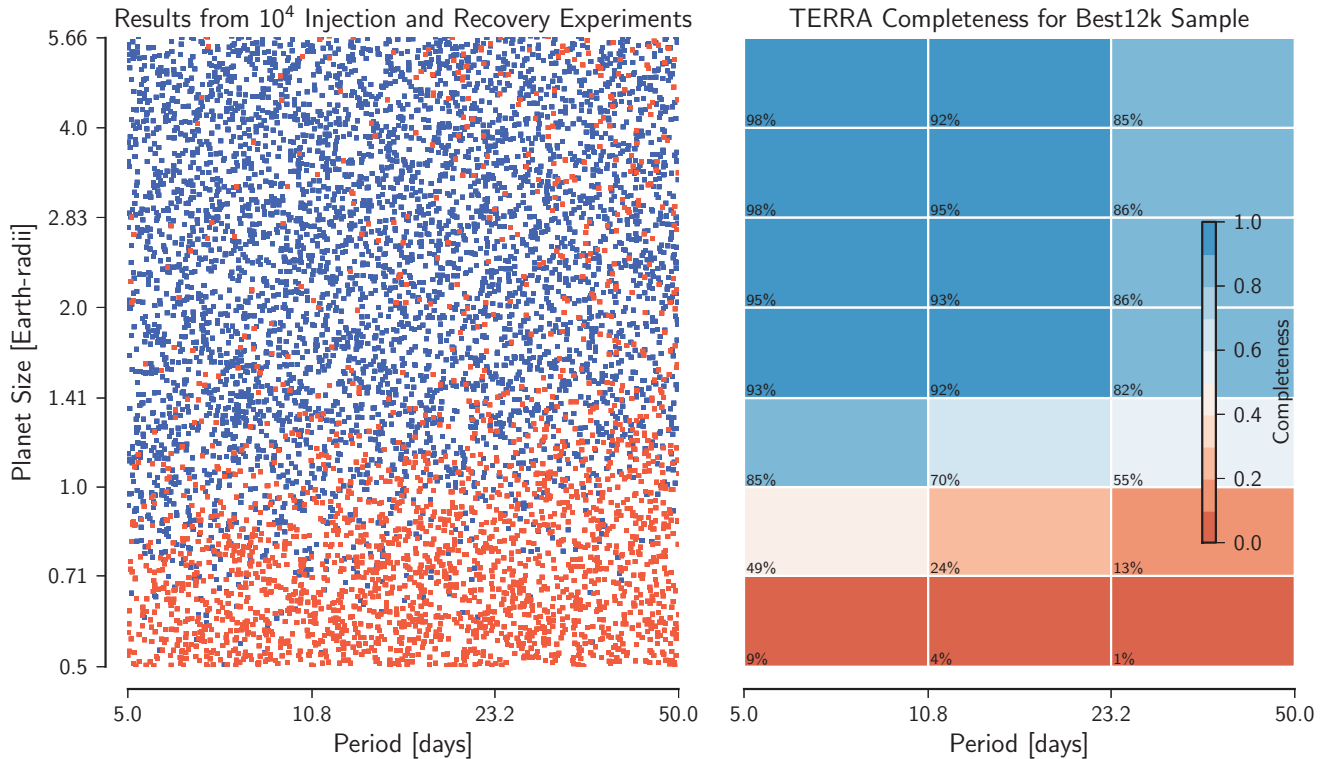


Figure 9. Results from the injection and recovery of 10,000 synthetic transit signals into actual photometry of randomly selected stars from our Best12k stellar sample. Each point represents the planet radius and orbital period of a mock transiting planet. The blue points represent signals that passed the DV post-analysis and where TERRA recovers the correct period and epoch. Signals that did not pass DV and/or were not successfully recovered, are shown as red points. Pipeline completeness is simply the number of blue points divided by the total number of points in each bin. The figure shows that for planet sizes above $1.0 R_E$, our pipeline discovers over 50% of the injected planets, and presumably accomplishes a similar success rate for actual transiting planets. The completeness for planets larger than $1 R_E$ is thus high enough to compute planet occurrence for such small planets, with only moderate completeness corrections needed (less than a factor of 2). Note that we are measuring the recovery rate of putative planets in the Best12k sample with TERRA. Had we selected a lower noise stellar sample, for example the “Best6k,” the region of high completeness would extend to even small radii.

tries. We then use Kepler’s 3rd law together with P and M_* to compute a/R_* assuming a circular orbit.⁵

To compute occurrence, we divide the number of stars with planets in a particular cell by the number of stars amenable to the detection of a planet in a given cell, $n_{*,\text{amen}}$. This number is just $N_* = 12,000$ times the completeness, computed in our Monte Carlo study. The debiased fraction of stars with planets per P - R_P bin, f_{cell} , is given by $f_{\text{cell}} = n_{\text{pl, aug, cell}}/n_{*,\text{amen}}$. We show f_{cell} on the P - R_P plane in Figure 10 as a color scale. We also compute $d^2 f_{\text{cell}}/d \log P/d \log R_P$, i.e. planet occurrence divided by the logarithmic area of each cell, which is a measure of occurrence which does not depend on bin size. We annotate each P - R_P bin of Figure 10 with the corresponding value of $n_{\text{pl, cell}}$, $n_{\text{pl, aug, cell}}$, f_{cell} , and $d^2 f_{\text{cell}}/d \log P/d \log R_P$.

Due to the small number of planets in each cell, errors due to counting statistics alone are significant. We compute Poisson errors on $n_{\text{pl, cell}}$ for each cell. Errors on $n_{\text{pl, aug, cell}}$, f_{cell} , and $d^2 f_{\text{cell}}/d \log P/d \log R_P$ include only the Poisson errors from $n_{\text{pl, cell}}$. There is also shot noise associated with the Monte Carlo completeness correction due to the finite number of simulated planets in each P - R_P cell, but such errors are small compared to errors on $n_{\text{pl, cell}}$. The orbital alignment correction, a/R_* , is also uncertain due to imperfect knowledge of stellar radii and orbital separations. We do not include such errors in our occurrence estimates.

Of particular interest is the distribution of planet occurrence with R_P for all periods. We marginalize over P by summing occurrence over all period bins from 5 to 50 days. The distribution of radii shown in Figure 11 shows a rapid rise in occurrence from 8.0 to 2.8 R_E . H12 also observed a rising occurrence of planets down to 2.0 R_E , which they modeled as a power law. *Planet occurrence is consistent with a flat distribution from 2.8 to 1.0 R_E , ruling out a continuation of a power law increase in occurrence for planets smaller than 2.0 R_E . We find $15.1^{+1.8}_{-2.7}\%$ of Sun-like stars harbor a 1.0–2.0 R_E planet with $P = 5$ –50 days. Including larger planets, we find that $24.8^{+2.1}_{-3.4}\%$ of stars harbor a planet larger than Earth with $P = 5$ –50 days. Occurrence values assuming a 100% efficient pipeline are shown as gray bars in Figure 11. The red bars show the magnitude of our completeness correction. Even though TERRA detects many planets smaller than 1.0 R_E , we do not report occurrence for planets smaller than Earth since pipeline completeness drops abruptly below 50%.*

We show planet occurrence as a function of orbital period in Figure 12. In computing this second marginal distribution, we include radii larger than 1 R_E so that corrections due to incompleteness are small. Again, as in Figure 11, gray bars represent uncorrected occurrence values while red bars show our correction to account for planets that TERRA missed. Planet occurrence rises as orbital period increases from 5.0 to 10.8 days. Above 10.8 days, planet occurrence is nearly constant per logarithmic period bin with a slight indication of a continued rise. This leveling off of the distribution was noted by H12, who considered $R_P > 2.0 R_E$. We fit the distribu-

⁵ H12 determined a/R_* directly from light curve fits, but found little change when computing occurrence from a/R_* using Kepler’s third law.

tion of orbital periods for $R_P > 1.0 R_E$ with two power laws of the form

$$\frac{df}{d \log P} = k_P P^\alpha, \quad (7)$$

where α and k_P are free parameters. We find best fit values of $k_P = 0.185^{+0.043}_{-0.035}$, $\alpha = 0.16 \pm 0.07$ for $P = 5$ –10.8 days and $k_P = 8.4^{+0.9}_{-0.8} \times 10^{-3}$, $\alpha = 1.35 \pm 0.05$ for $P = 10.8$ –50 days. We note that k_P and α are strongly covariant. Extrapolating the latter fit speculatively to $P > 50$ days, we find $41.7^{+6.8}_{-5.9}\%$ of Sun-like stars host a planet 1 R_E or larger with $P = 50$ –500 days.

7. COMPARISON OF TERRA AND BATALHA ET AL. (2012) PLANET CATALOGS

Here, we compare our candidates to those of Batalha et al. (2012). Candidates were deemed in common if their periods agree to within 0.01 days. We list the union of the TERRA and Batalha et al. (2012) catalogs in Table 4. Eighty-two candidates appear in both catalogs (Section 7.1), 47 appear in this work only (Section 7.2), and 33 appear in Batalha et al. (2012) only (Section 7.3). We discuss the significant overlap between the two catalogs and explain why some candidates were detected by one pipeline but not the other.

7.1. Candidates in Common

Eighty-two of our candidates appear in the Batalha et al. (2012) catalog. We show these candidates in P - R_P space in Figure 8 as grey points. TERRA detected no new candidates with $R_P > 2 R_E$. This agreement in detected planets having $R_P > 2 R_E$ demonstrates high completeness for such planets in both pipelines for this sample of quiet stars. This is not very surprising since candidates with $R_P > 2 R_E$ have high SNR, e.g. min, median, and max SNR = 19.3, 71.5, and 435 respectively.

Radii for the 82 planets in common were fairly consistent between Batalha et al. (2012) and this work. The two exceptions were KIC-8242434 and KIC-8631504. Using SpecMatch, we find stellar radii of 0.68 and 0.72 R_\odot , respectively, down from 1.86 and 1.80 R_\odot in Batalha et al. (2012). The revised planet radii are smaller by over a factor of two. Radii for the other planets in common were consistent to $\sim 20\%$.

7.2. TERRA Candidates Not in Batalha et al. (2012) Catalog

TERRA revealed 47 planet candidates that did not appear in Batalha et al. (2012). Such candidates are colored blue and red in Figure 8. Many of these new detections likely stem from the fact that we use twice the photometry that was available to Batalha et al. (2012). To get a sense of how additional photometry improves the planet yield of the *Kepler* pipeline beyond Batalha et al. (2012), we compared the TERRA candidates to the *Kepler* team KOI list dated August 8, 2012 (Jason Rowe, private communication). The 28 candidates in common between the August 8, 2012 *Kepler* team sample and this work are colored blue in Figure 8. Of these 28 candidates, 10 are listed as false positives and denoted as crosses in Figure 8.

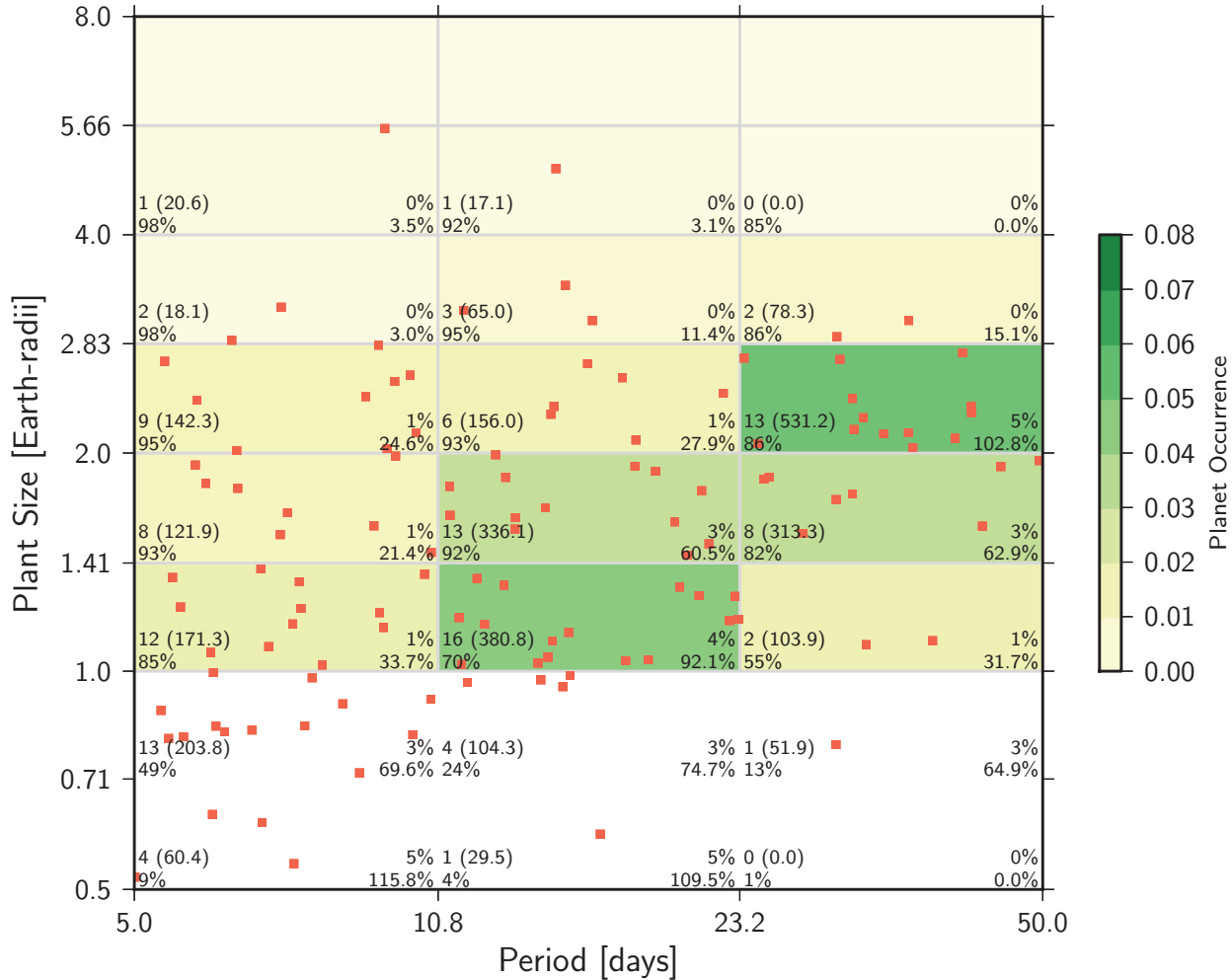


Figure 10. Planet occurrence as a function of orbital period and planet radius for $P = 5\text{--}50$ days and $R_P = 0.5\text{--}8 R_E$. TERRA planet candidates are shown as red points. Cell occurrence, f_{cell} , is given by the color scale. We quote the following information for each cell: Top left—number of planets (number of augmented planets); lower left—completeness; top right—fractional planet occurrence, f_{cell} ; bottom right—normalized planet occurrence, $d^2 f_{\text{cell}} / d \log P / d \log R_P$. We do not color cells where the completeness is less than 50% (i.e. the completeness correction is larger than a factor of 2).

We announce 37 new planet candidates with respect to Batalha et al. (2012) that were not listed as false positives in the *Kepler* team sample. These 37 candidates, all with $R_P \lesssim 2 R_E$, are a subset of those listed in Table 2. As a convenience, we show this subset in Table 3. We remind the reader that all photometry used in this work is publicly available. We hope that interested readers will fold the photometry on the ephemeris in Table 2 and access critically whether a planet interpretation is correct. As a quick reference, we have included plots of the transits of the 37 new candidates from Table 3 in the appendix (Figures 17 and 18). We do not claim that our additional candidates bring pipeline completeness to unity for planets with $R_P \lesssim 2 R_E$. As shown in Section 5, our planet sample suffers from significant incompleteness in the same P - R_P space where most of the new candidates emerged.

7.3. Batalha et al. (2012) Candidates Not in TERRA catalog

There are 33 planet candidates in the Batalha et al. (2012) catalog from Best12k stars that TERRA missed. Of these, 28 are multi-candidate systems where one component was identified by TERRA. TERRA is currently insensitive to multiple planet systems (as described in Section 3.2). TERRA missed the remaining 5 Batalha et al. (2012) candidates for the following reasons:

- 2581.01 : A bug in the pipeline prevented successful photometric calibration (Section 3.1). This bug affected 19 out of 12,000 stars in the Best12k sample.
- 70.01, 111.01, 119.01 : Failed one of the automated DV cuts (`taur`, `med_on_mean`, and `taur`, respectively). We examined these three light curves in the fashion described in Section 3.3, and we determined these light curves were consistent with an exoplanet transit. The fact that DV is discarding compelling transit signals decreases TERRA’s overall completeness. Computing DV metrics and choosing the optimum cuts is an art. There is room for

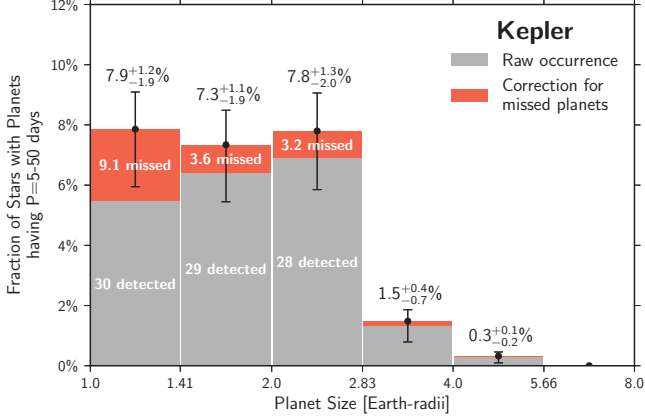


Figure 11. Distribution of planet occurrence for R_P ranging from 1.0 to 8.0 R_E . We quote the fraction of Sun-like stars harboring a planet with $P = 5\text{--}50$ days for each R_P bin. We observe a rapid rise in planet occurrence from 8.0 down to 2.8 R_E , as seen in H12. Below 2.8 R_E , the occurrence distribution is consistent with flat. This result rules out a power law increase in planet occurrence toward smaller radii. Adding up the two smallest radius bins, we find $15.1^{+1.8}_{-2.7}\%$ of Sun-like stars harbor a 1.0–2.0 R_E planet within ~ 0.25 AU. To compute occurrence as a function of R_P , we simply sum occurrence rates for all period bins shown in Figure 10. Errors due to counting statistics are computed by adding errors from each of the three period bins in quadrature. The gray portion of the histogram shows occurrence values before correcting for missed planets due to pipeline incompleteness. Our correction to account for missed planets is shown in red, and is determined by the injection and recovery of synthetic transits described in Section 5. We do not show occurrence values where the completeness is $< 50\%$.

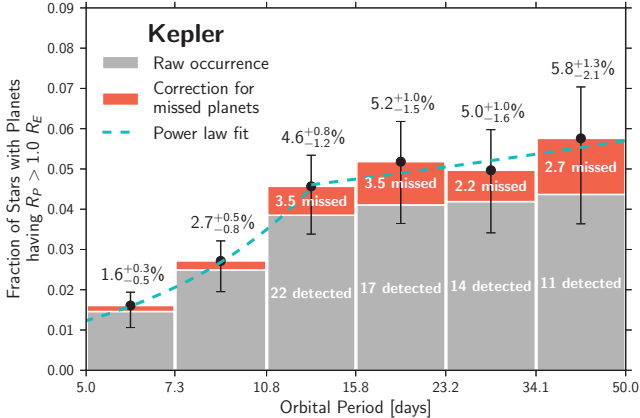


Figure 12. Distribution of planet occurrence for different orbital periods ranging from 5 to 50 days. We quote the fraction of Sun-like stars with a planet Earth-size or larger as a function of orbital period. We observe a gradual rise in occurrence from 5.0 to 10.8 R_E followed by a leveling off for longer orbital periods. H12 observed a similar leveling off in their analysis which included planets larger than 2 R_E . We fit the domains above and below 10.8 R_E separately with power laws, $df/d\log P = k_P P^\alpha$. We find best fit values of $k_P = 0.185^{+0.043}_{-0.035}$, $\alpha = 0.16 \pm 0.07$ for $P = 5\text{--}10.8$ days and $k_P = 8.4^{+0.9}_{-0.8} \times 10^{-3}$, $\alpha = 1.35 \pm 0.05$ for $P = 10.8\text{--}50$ days. Speculatively, we extrapolate the latter power law fit another decade in period and estimate $41.7^{+6.8}_{-5.9}\%$ of Sun-like stars harbor a planet Earth-size or larger with $P = 50\text{--}500$ days. As in Figure 11, the gray portion of the histogram shows uncorrected occurrence while the red region shows our correction for pipeline incompleteness. Note that the number of detected planets decreases as P increases from 10.8 to 50 days, while occurrence remains nearly constant. At longer periods, the geometric transit probability is lower, and each detected planet counts more toward $df/d\log P$.

improvement here.

- KOI-1151.01 : Period misidentified in Batalha et al. (2012). In Batalha et al. (2012) KOI-1151.01 is listed with with a $P = 5.22$ days. TERRA found a candidate with $P = 10.43$ days. Figure 13 shows phase-folded photometry with the TERRA ephemeris. A period of 5.22 days would imply dimmings in regions where the light curve is flat.

We plot the 33 total candidates listed in Batalha et al. (2012), but not found by TERRA in Figure 14. We highlight the 5 missed candidates that cannot be explained by the fact that they are a lower SNR candidate in a multi-candidate system. TERRA is blind to planets in systems with another planet with higher SNR. Figure 14 shows that most of these missed planets occur at $R_P < 1.4 R_E$.

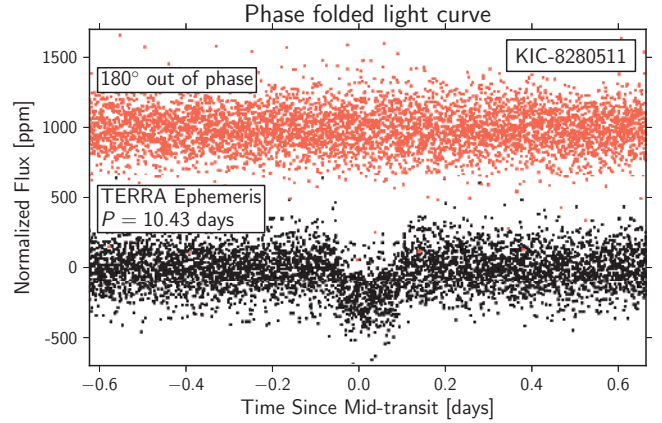


Figure 13. Phase-folded photometry of KIC-8280511 folded on the correct 10.43 day period found by TERRA. KOI-1151.01 is listed with $P = 5.22$ days in Batalha et al. (2012). If the transit was truly on the 5.22 day period, we should see a transit of equal depth 180 degrees out of phase. KOI-1151.01 is listed in Batalha et al. (2012) with half its true period.

8. OCCURRENCE WITH PLANET MULTIPLICITY INCLUDED

While the TERRA planet occurrence measurement benefits from well-characterized completeness, it does not include the contribution of multis to overall planet occurrence. As discussed in Sections 3.2 and 7.3, TERRA only detects the highest SNR candidate for a given star. Here, we present planet occurrence including multis from Batalha et al. (2012). Thus, the occurrence within a bin, f , in this section should be interpreted as the *average number of planets* per star with $P = 5\text{--}50$ days. The additional planets from Batalha et al. (2012) raise the occurrence values somewhat over those of the previous section. However, the rise and plateau structure remains the same.

We compute f_{cell} from the 32 candidates present in Batalha et al. (2012), but not found by TERRA (mis-labeled KOI-1151.01 was not included). For clarity, we refer to this separate occurrence calculation as $f_{\text{cell},\text{Batalha}}$. Because the completeness of the *Kepler* pipeline is unknown, we apply no completeness correction. This assumption of 100% completeness is certainly an overestimate, but we believe that the sensitivity of the *Kepler*

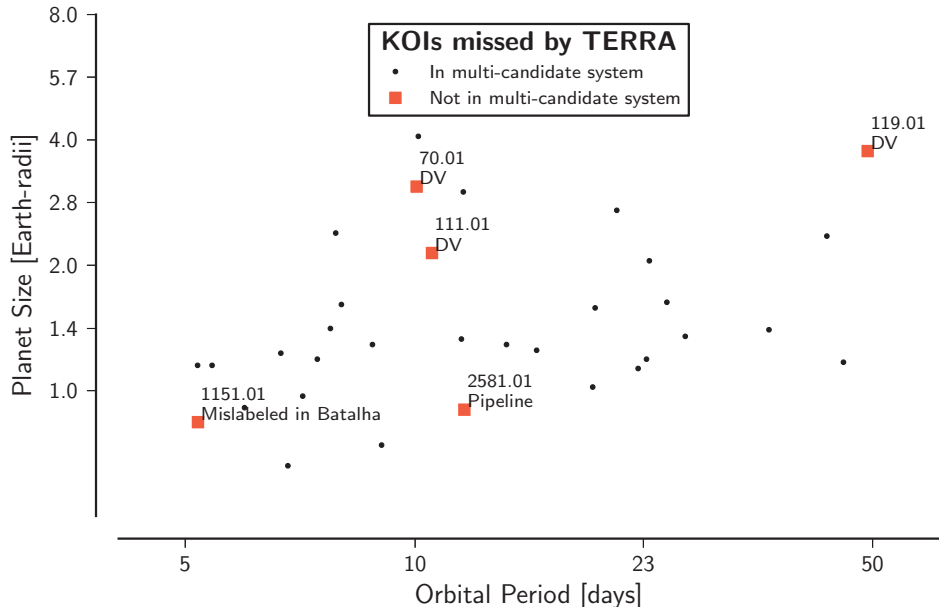


Figure 14. P and R_P for the 33 candidates present in Batalha et al. (2012) but not found by TERRA. The small symbols show the candidates in multi-planet systems. TERRA is blind to such candidates. The 5 larger symbols show the other failure modes of TERRA: 2581.01 failed due to a pipeline bug; 70.01, 111.01, and 119.01 did not pass DV; and TERRA missed KOI-1151.01 because it is listed in Batalha et al. (2012) with the incorrect period. Most of the missed planets have $R_P < 1.4 R_E$.

pipeline to multis is nearly complete for $R_P > 1.4 R_E$. TERRA has $> 80\%$ completeness for $R_P > 1.4 R_E$ because planets in that size range with $P = 5\text{--}50$ days around Best12k stars have high SNR. The *Kepler* pipeline should also be detecting these high SNR candidates. Also, once a KOI is found, the *Kepler* team reprocesses the light curve for additional transits (Jason Rowe, private communication). Due to this additional scrutiny, we believe that the *Kepler* completeness for multis is higher than for singles, all else being equal.

We then add $f_{\text{cell}, \text{Batalha}}$ to f_{cell} computed in the previous section. We show occurrence computed using TERRA and Batalha et al. (2012) planets as a function of P and R_P in Figure 15 and as a function of only R_P in Figure 16. The 32 additional planets from Batalha et al. (2012) do not change the overall shape of the occurrence distribution: rising from 4.0 to 2.8 R_E and consistent with flat from 2.8 down to 1.0 R_E .

H12 fit occurrence for $R_P > 2 R_E$ with a power law,

$$\frac{df}{d \log R_P} = k_R R_P^\alpha, \quad (8)$$

finding $\alpha = -1.92 \pm 0.11$ and $k_R = 2.9^{+0.5}_{-0.4}$ (Section 3.1 of H12). As a point of comparison, we plot the H12 power law over our combined occurrence distribution in Figure 16. The fit agrees qualitatively for $R_P > 2 R_E$, but not within errors. We expect the H12 fit to be $\sim 25\%$ higher than our occurrence measurements since H12 included planets with $P < 50$ days (not $P = 5\text{--}50$ days). Additional discrepancies could stem from different characterizations of completeness, reliance on photometric versus spectroscopic measurements of R_* , and magnitude-limited, rather than noise-limited, samples.

9. DISCUSSION

9.1. TERRA

We implement in this work a new pipeline for the detection of transiting planets in *Kepler* photometry and apply it to a sample of 12,000 G and K-type dwarfs stars chosen to be among the most photometrically quiet of the *Kepler* target stars. These low noise stars offer the best chance for the detection of small, Earth-size planets in the *Kepler* field and will one day be among the stars from which η_\oplus —the fraction of Sun-like stars bearing Earth-size planets in habitable zone orbits—is estimated. In this work, we focus on the close-in planets having orbital periods of 5–50 days and semi-major axes $\lesssim 0.25$ AU. Earth-size planets with these characteristics are statistically at the margins of detectability with the current ~ 3 years of photometry in *Kepler* quarters Q1–Q12.

Our TERRA pipeline has two key features that enable confident measurement of the occurrence of close-in planets approaching Earth size. First, TERRA calibrates the *Kepler* photometry and searches for transit signals independent of the results from the *Kepler* Mission’s official pipeline. In some cases, TERRA calibration achieves superior noise suppression compared to the Pre-search Data Conditioning (PDC) module of the official *Kepler* pipeline (Petigura & Marcy 2012). The transit search algorithm in TERRA is efficient at detecting low SNR transits in the calibrated light curves. This algorithm successfully rediscovers 82 of 86 stars bearing planets in Batalha et al. (2012). Recall that the current version of TERRA only detects the highest SNR transit signal in each system. Thus, additional planets orbiting known hosts are not reported here. We report the occurrence of stars having one or more planets, not the mean number of planets per star as in H12 and elsewhere. Our pipeline also detects 37 planets not found in Batalha et al. (2012) (19 of which were not in the catalog of the *Kepler* team as of August 8th, 2012), albeit with the benefit of 6 quarters of additional photometry for TERRA to search.

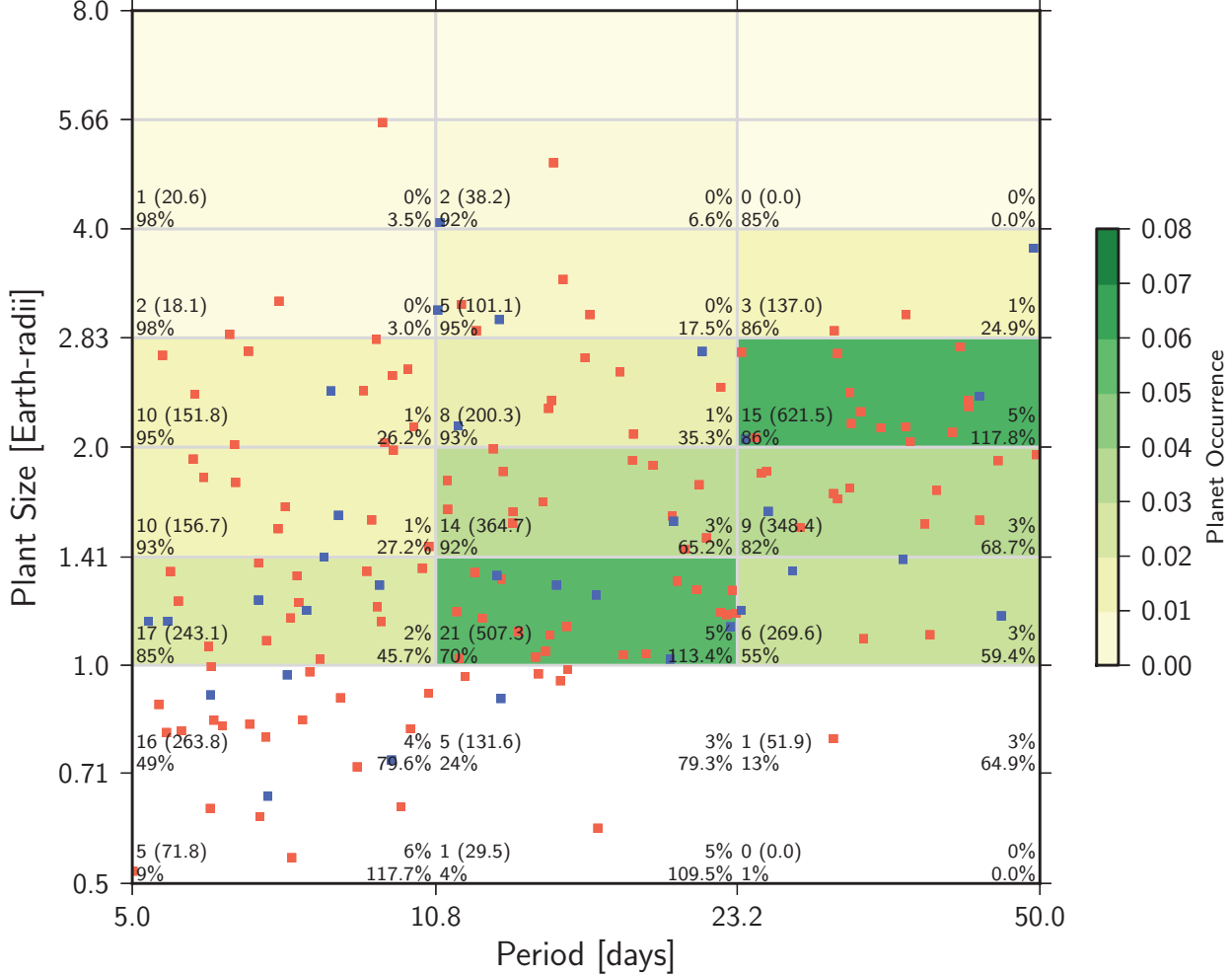


Figure 15. As in Figure 10, red points show 119 **TERRA**-detected planets. Blue points represent additional planets from Batalha et al. (2012). Most (28 out of 32) of these new candidates are planets in multi-candidate systems where **TERRA** successfully identifies the higher SNR candidate. We apply no completeness correction to these new planets, and we believe this is appropriate for $R_P > 1.4 R_E$. We quote the following occurrence information for each cell: Top left—number of planets (number of augmented planets), lower left—completeness, top right—fractional planet occurrence f_{cell} , bottom right—normalized planet occurrence $d^2 f_{\text{cell}}/d \log P/d \log R_P$. We do not color cells where the completeness is less than 50% (i.e. the completeness correction is larger than a factor of 2). The planet counts and occurrence values are for the combined **TERRA** and Batalha et al. (2012) sample. The completeness values are the same as in Figure 10.

The second crucial feature of **TERRA** is that we have characterized its detection completeness via the injection and recovery of synthetic transits in real *Kepler* light curves from the Best12k sample. *This completeness study is crucial to our occurrence calculations because it allows us to statistically correct for incompleteness variations across the P - R_P plane.* While the *Kepler* Project has initiated a completeness study of the official pipeline (Christiansen et al. 2012), **TERRA** is the only pipeline for *Kepler* photometry whose detection completeness has been calibrated by injection and recovery tests. Prior to **TERRA**, occurrence calculations required one to *assume* that the *Kepler* planet detections were complete down to some SNR limit, or to estimate completeness based on SNR alone without empirical tests of the performance of the algorithms in the pipeline. For example, H12 made cuts in stellar brightness ($Kp < 15$) and transit SNR (> 10 in a single quarter of photometry) and restricted their search to planets larger than $2 R_E$ with orbital periods shorter than 50 days. These

conservative cuts on the planet and star catalogs were driven by the unknown completeness of the official *Kepler* pipeline at low SNR. H12 applied two statistical corrections to convert their distribution of detected planets into an occurrence distribution. They corrected for non-transiting planets with a geometric a/R_* correction. They also computed the number of stars amenable to the detection (at $\text{SNR} > 10$ in a single quarter) of each planet and considered only that number of stars in the occurrence calculation. H12 had no empirical way to determine the actual detection efficiency of the algorithms in the pipeline. Here, we apply the geometric a/R_* correction and correct for pipeline completeness across the P - R_P plane by explicit tests of the **TERRA** pipeline efficiency, which naturally incorporates an SNR threshold correction as in H12.

9.2. Planet Occurrence

H12 found that for close-in planets, the planet radius function rises steeply from Jupiter size to $2 R_E$. For

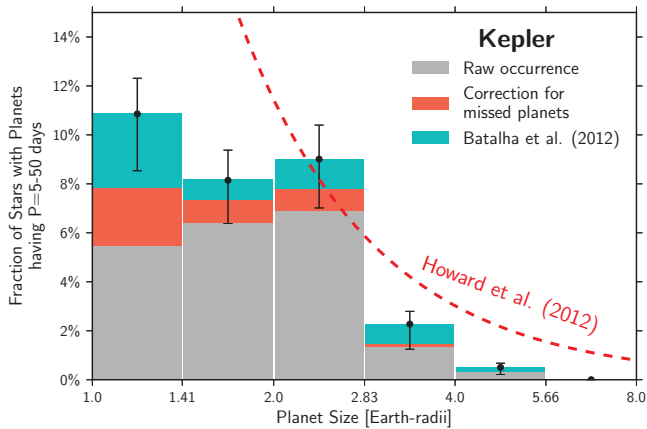


Figure 16. Same as Figure 11 with inclusion of planets in multi planet systems. The blue regions represent the additional contribution to planet occurrence from the Batalha et al. (2012) planets. The addition of these new planets does not change the overall shape of the distribution. The dashed line is the power law fit to the planet size distribution in H12. The fit agrees qualitatively for $R_P > 2 R_E$, but not within errors. We expect the H12 fit to be $\sim 25\%$ higher than our occurrence measurements since H12 included planets with $P < 50$ days (not $P = 5-50$ days). Additional discrepancies between occurrence in H12 and this work could stem from different characterizations of completeness, reliance on photometric versus spectroscopic measurements of R_* , and magnitude-limited, rather than noise-limited, samples.

smaller planets of $\sim 1-2 R_E$, occurrence was approximately constant in logarithmic R_P bins, but H12 were skeptical of the result below $2 R_E$ because of unknown pipeline completeness and the small number statistics near $1 R_E$ in the Borucki et al. (2011b) planet catalog. In this work, we strongly confirm the power law rise in occurrence from 4 to $2 R_E$ using a superior assessment of completeness and nine times more photometry than in H12. Using TERRA, we can empirically and confidently compute occurrence down to $1 R_E$. Our key result is the plateau of planet occurrence for the size range $1-2.8 R_E$ for planets having orbital periods 5–50 days around Sun-like stars. In that size range of $1-2.8 R_E$, 23% of stars have a planet orbiting with periods between 5 and 50 days. Including the multiple planets within each system, we find 0.28 planets per star within the size range $1-2.8 R_E$ and with periods between 5–50 days. These results apply, of course, to the *Kepler* field, with its still unknown distribution of masses, ages, and metallicities in the Galactic disk.

As shown in Figure 10, TERRA detects many sub-Earth size planets ($< 1.0 R_E$). These sub-Earths appear in regions of low completeness, and, provocatively, may represent just the tip of the iceberg. A rich population of sub-Earths may await discovery given more photometry and continued pipeline improvements. With 8 years of total photometry in an extended *Kepler* mission (compared to 3 years here), the computational machinery of TERRA—including its light curve calibration, transit search, and completeness calibration—will enable a measurement of η_{\oplus} for habitable zone orbits.

9.3. Interpretation

We are not the first to note the huge population of close-in planets smaller or less massive than Neptune. Using Doppler surveys, Howard et al. (2010) and Mayor et al. (2011) showed that the planet mass func-

tion rises steeply with decreasing mass, at least for close-in planets. In *Kepler* data, the excess of close-in, small planets was obvious in the initial planet catalogs released by the *Kepler* Project (Borucki et al. 2011a,b). H12 characterized the occurrence distribution of these small planets as a function of their size, orbital period, and host star temperature. These occurrence measurements, based on official *Kepler* planet catalogs, were refined and extended by Youdin (2011), Traub (2012), Dong & Zhu (2012), Beaugé & Nesvorný (2013), and others. Our contribution here shows a clear plateau in occurrence in the $1-2.8 R_E$ size range and certified by an independent search of *Kepler* photometry using a pipeline calibrated by injection and recovery tests. The onset of the plateau at $\sim 2.8 R_E$ suggests that there is a preferred size scale for the formation of close-in planets.

H12 and Youdin (2011) noted falling planet occurrence for periods shorter than ~ 7 days. We also observe declining planet occurrence for short orbital periods, but find that the transition occurs closer to ~ 10 days. We consider our period distribution to be in qualitative agreement with those of H12 and Youdin (2011). Planet formation and/or migration seems to discourage very close-in planets ($P \lesssim 10$ days).

Close-in, small planets are now the most abundant planets detected by current transit and Doppler searches, yet they are absent from the solar system. The solar system is devoid of planets between 1 and $3.88 R_E$ (Earth and Neptune) and planets with periods less than Mercury’s ($P < 88.0$ days). The formation mechanisms and possible subsequent migration of such planets are hotly debated. The population synthesis models of Ida & Lin (2010) and Mordasini et al. (2012) suggest that they form near or beyond the ice line and then migrate quiescently in the protoplanetary disk. These models follow the growth and migration of planets over a wide range of parameters (from Jupiter mass down to Earth mass orbiting at distances out to ~ 10 AU) and they predict “deserts” of planet occurrence that are not detected.

More recently, Hansen & Murray (2012) and Chiang & Laughlin (2012) have argued for the *in situ* formation of close-in planets of Neptune size and smaller. In these models, close-in rocky planets of a few Earth masses form from protoplanetary disks more massive than the minimum mass solar nebula. Multiple planets per disk form commonly in these models and accretion is fast ($\sim 10^5$ years) and efficient due to the short dynamical timescales of close-in orbits. The rocky cores form before the protoplanetary disk has dissipated, accreting nebular gas that adds typically $\sim 3\%$ to the mass of the planet (Chiang & Laughlin 2012). But the small amounts of gas can significantly swell the radii of these otherwise rocky planets. For example, Adams et al. (2008) found that adding a H/He gas envelope equivalent to 0.2–20% of the mass of a solid $5 M_E$ planet increases the radius 8–110% above the gas-free value.

We find the *in situ* model plausible because it naturally explains the large number of close, sub-Neptune-size planets, the high rate of planet multiplicity and nearly co-planar and circular orbits (Lissauer et al. 2011; Fang & Margot 2012), and does not require tuning of planet migration models. Our result of a plateau in the planet size distribution for $1-2.8 R_E$ with a sharp falloff

in occurrence for larger planets along with decreasing occurrence for $P \lesssim 10$ days are two significant observed properties of planets around Sun-like stars that must be reproduced by models that form planets *in situ* or otherwise and by associated population synthesis models.

The *in situ* model seems supported by the sheer large occurrence of sub-Neptune-size planets within 0.25 AU. It seems unlikely that all such planets form beyond the snow line at ~ 2 AU, which would require inward migration to within 0.25 AU, but not all the way into the star. Such models of formation beyond the snow line seem to require fine tuning of migration and parking mechanisms, as well as the tuning of available water or gas beyond 2 AU, while avoiding runaway gas accretion toward Jupiter masses. Still, *in situ* formation seems to require higher densities than those normally assumed in a minimum mass solar nebula (Chiang & Laughlin 2012) in order to form the sub-Neptune planets before removal of the gas. If this *in situ* model is correct, we expect these sub-Neptune-size planets to be composed of rock plus H and He, rather than rock plus water (Chiang & Laughlin 2012). Thus, a test of the *in situ* mode of formation involves spectroscopic measurements of the chemical composition of the close-in sub-Neptunes.

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REFERENCES

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- 08, 1
 Adams, E. R., Seager, S., & Elkins-Tanton, L. 2008, ApJ, 673, 1160
 Batalha, N. M., et al. 2012, ArXiv e-prints
 Beaugé, C., & Nesvorný, D. 2013, ApJ, 763, 12
 Borucki, W. J., et al. 2011a, ApJ, 728, 117
 —. 2011b, ApJ, 736, 19
 Brown, T. M., Latham, D. W., Everett, M. E., & Esquerdo, G. A. 2011, AJ, 142, 112
 Chiang, E., & Laughlin, G. 2012, ArXiv e-prints
 Christiansen, J., Science Office, K., & Science Operations Center, K. 2012, in AAS/Division for Planetary Sciences Meeting Abstracts, Vol. 44, AAS/Division for Planetary Sciences Meeting Abstracts, 113.01
 Collette, A. 2008, HDF5 for Python
 Demarque, P., Woo, J.-H., Kim, Y.-C., & Yi, S. K. 2004, ApJS, 155, 667
 Dong, S., & Zhu, Z. 2012, ArXiv e-prints
 Fang, J., & Margot, J.-L. 2012, ApJ, 761, 92
 Fraquelli, D., & Thompson, S. E. 2012, Kepler Archive Manual (KDMC-10008-004)
 Fressin, F., et al. 2013, ArXiv e-prints
 Hansen, B. M. S., & Murray, N. 2012, ApJ, 751, 158
 Howard, A. W., et al. 2010, Science, 330, 653
 —. 2012, ApJS, 201, 15
 Hunter, J. D. 2007, Computing In Science & Engineering, 9, 90
 Ida, S., & Lin, D. N. C. 2010, ApJ, 719, 810
 Jenkins, J. M., et al. 2010a, ApJ, 713, L87
 Jenkins, J. M., et al. 2010b, in Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, Vol. 7740, Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series
 Jones, E., Oliphant, T., Peterson, P., et al. 2001–, SciPy: Open source scientific tools for Python
 Kovács, G., Zucker, S., & Mazeh, T. 2002, A&A, 391, 369
 Lissauer, J. J., et al. 2011, ApJS, 197, 8
 Mandel, K., & Agol, E. 2002, ApJ, 580, L171
 Marcy, G. W., et al. 2008, Physica Scripta Volume T, 130, 014001
 Mayor, M., et al. 2011, arXiv:1109.2497
 Mordasini, C., Alibert, Y., Klahr, H., & Henning, T. 2012, A&A, 547, A111
 Muirhead, P. S., et al. 2012, ApJ, 747, 144
 Oliphant, T. E. 2007, Computing in Science Engineering, 9, 10
 Pérez, F., & Granger, B. E. 2007, Comput. Sci. Eng., 9, 21
 Petigura, E. A., & Marcy, G. W. 2012, PASP, 124, 1073
 Smith, J. C., et al. 2012, PASP, 124, 1000
 Staelin, D. H. 1969, in Proceedings of the IEEE, National Radio Astronomy Observatory, Charlottesville, VA, USA, 724–725
 Stumpe, M. C., et al. 2012, PASP, 124, 985
 Traub, W. A. 2012, ApJ, 745, 20
 Twicken, J. D., Chandrasekaran, H., Jenkins, J. M., Gunter, J. P., Girouard, F., & Klaus, T. C. 2010, in Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, Vol. 7740, Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series
 Youdin, A. N. 2011, ApJ, 742, 38

APPENDIX

Table 2
 Planet candidates identified with TERRA

kic	Light Curve Fit							Stellar Parameters								
	P (d)	t_0^a (d)	$\frac{R_P}{R_\star}$ (%)	$\sigma(\frac{R_P}{R_\star})$	τ (hrs)	$\sigma(\tau)$	b	$\sigma(b)$	T_{eff} (K)	$\log g$ (cgs)	R_\star (R_\odot)	R_P (R_E)	$\sigma(R_P)$	source ^b	FP	B12
2142522	13.323	67.043	0.98	0.34	2.32	0.51	< 0.85		6046	4.40	1.04	1.11	0.39	P1	Y	N
2307415	13.122	66.396	1.26	0.13	1.90	0.12	< 0.52		6133	4.38	1.14	1.57	0.16	S	N	Y
2441495	12.493	71.457	2.38	0.24	1.13	0.03	< 0.47		5192	4.56	0.76	1.99	0.20	S	N	Y
2444412	14.911	74.337	3.38	0.34	4.24	0.31	0.93	0.01	5551	4.47	0.92	3.41	0.34	S	N	Y
2571238	9.287	68.984	2.90	0.29	3.72	0.38	0.90	0.03	5544	4.50	0.89	2.82	0.28	S	N	Y
2853446	7.373	70.613	1.39	0.14	0.74	0.06	< 0.70		5969	4.37	1.10	1.65	0.17	S	N	Y
3098810	40.811	75.673	1.98	0.20	1.45	0.12	< 0.70		6071	4.31	1.27	2.75	0.28	S	N	Y

Table 2 — Continued

kic	Light Curve Fit								Stellar Parameters							
	P (d)	t_0^a (d)	$\frac{R_P}{R_\star}$ (%)	$\sigma(\frac{R_P}{R_\star})$	τ (hrs)	$\sigma(\tau)$	b	$\sigma(b)$	T_{eff} (K)	$\log g$ (cgs)	R_\star (R_\odot)	R_P (R_E)	$\sigma(R_P)$	source ^b	FP	B12
3120904	42.915	72.866	1.15	0.12	2.96	0.36	< 0.66		6151	4.31	1.26	1.59	0.16	S		N
3342794	14.172	75.591	1.40	0.14	1.31	0.13	< 0.55		5900	4.35	1.10	1.68	0.17	S	N	Y
3442055	29.619	66.681	1.57	0.16	2.48	0.15	< 0.57		5624	4.41	1.01	1.72	0.17	S	N	Y
3531558	24.994	71.674	1.49	0.15	2.93	0.12	< 0.53		5808	4.35	1.14	1.85	0.19	S	N	Y
3545135	8.483	65.973	0.81	0.08	1.44	0.10	< 0.57		5794	4.40	1.02	0.90	0.09	S	N	N
3835670	14.558	78.084	2.86	0.29	3.89	0.22	0.31	0.19	5722	4.14	1.58	4.93	0.49	S	N	Y
3839488	11.131	67.370	1.36	0.14	1.97	0.11	< 0.50		5991	4.36	1.11	1.64	0.17	S	N	Y
3852655	11.629	65.817	0.85	0.30	1.49	0.23	< 0.79		5822	4.36	1.05	0.97	0.34	P1	N	N
3942670	33.416	70.911	1.48	0.15	3.99	0.24	< 0.63		6012	4.28	1.32	2.13	0.21	S	N	Y
4043190	6.401	69.883	1.07	0.11	1.35	0.08	< 0.65		5302	3.83	2.45	2.86	0.29	S	N	Y
4049901	16.291	65.115	0.64	0.23	1.72	0.16	< 0.68		5250	4.48	0.85	0.60	0.21	P1	N	Y
4548011	6.284	66.198	0.60	0.06	2.23	0.64	< 0.49		5991	4.30	1.26	0.83	0.09	S		N
4644604	14.486	64.550	2.03	0.21	1.49	0.15	< 0.59		5739	4.34	1.04	2.32	0.23	S	N	Y
4770174	6.096	67.600	0.57	0.20	3.08	0.44	< 0.75		6013	4.44	1.01	0.63	0.22	P1	N	N
4827723	7.239	68.024	1.63	0.17	2.20	0.37	0.70	0.12	5392	4.52	0.87	1.54	0.16	S	N	Y
4914423	15.965	75.182	2.17	0.22	3.14	0.28	0.71	0.06	5904	4.27	1.29	3.05	0.31	S	N	Y
4914566	22.241	77.736	0.83	0.29	3.76	0.36	< 0.73		5974	4.22	1.31	1.18	0.41	P1	Y	N
5009743	41.699	102.563	1.94	0.20	2.87	0.30	< 0.70		5937	4.35	1.09	2.32	0.23	S	N	Y
5042210	12.147	64.535	0.81	0.08	3.22	0.21	< 0.56		6007	4.27	1.31	1.16	0.12	S	N	Y
5094751	6.482	68.943	1.69	0.17	2.51	0.31	0.69	0.13	5929	4.37	1.10	2.02	0.20	S	N	Y
5096590	29.610	70.332	1.00	0.35	2.60	0.20	< 0.67		5623	4.63	0.73	0.79	0.28	P1		N
5121511	30.996	93.790	2.35	0.24	1.40	0.06	< 0.47		5217	4.53	0.84	2.15	0.22	S	N	Y
5308537	14.265	76.217	0.76	0.08	1.92	0.53	< 0.87		5831	4.29	1.26	1.05	0.11	S		N
5561278	20.310	79.826	1.16	0.12	2.73	0.15	< 0.47		6161	4.38	1.14	1.45	0.15	S	N	Y
5613330	23.449	69.031	1.68	0.17	4.44	0.28	0.25	0.22	6080	4.20	1.48	2.70	0.27	S	N	Y
5652893	14.010	65.959	1.15	0.13	1.98	0.49	< 0.85		5150	4.55	0.78	0.97	0.11	S	N	N
5702939	18.398	77.934	1.09	0.38	2.45	0.27	< 0.54		5634	4.47	0.87	1.04	0.36	P1	N	N
5735762	9.674	68.009	2.73	0.27	1.61	0.06	< 0.47		5195	4.53	0.84	2.51	0.25	S	N	Y
5866724	5.860	65.040	1.65	0.17	2.01	0.05	0.22	0.14	6109	4.29	1.31	2.37	0.24	S	N	Y
5959719	6.738	66.327	0.99	0.10	1.30	0.21	< 0.76		5166	4.56	0.77	0.83	0.09	S	N	Y
6071903	24.308	87.054	2.28	0.23	1.29	0.05	< 0.27		5296	4.55	0.83	2.06	0.21	S	N	Y
6197215	10.613	68.691	1.23	0.13	0.48	0.06	< 0.80		5933	4.39	1.09	1.46	0.15	S	N	N
6289257	19.675	69.903	1.33	0.13	1.96	0.21	0.35	0.27	6023	4.36	1.11	1.61	0.16	S	N	Y
6291837	35.596	84.942	2.45	0.25	3.18	0.42	< 0.71		6165	4.38	1.14	3.05	0.31	S	N	Y
6356692	11.392	74.555	0.66	0.23	2.74	0.31	< 0.73		5420	4.03	1.64	1.19	0.42	P1	N	N
6365156	10.214	73.050	1.57	0.16	2.83	0.09	< 0.42		5852	4.28	1.25	2.13	0.21	S	N	Y
6442340	13.137	76.973	1.38	0.14	2.41	0.13	< 0.55		5764	4.38	1.08	1.63	0.16	S	N	Y
6521045	12.816	68.772	1.37	0.14	3.42	0.34	0.40	0.19	5874	4.29	1.24	1.85	0.19	S	N	Y
6523351	6.067	69.128	0.72	0.25	1.02	0.20	< 0.89		5489	4.15	1.35	1.06	0.37	P1		N
6605493	9.310	69.379	0.99	0.10	1.80	0.16	< 0.62		5805	4.36	1.12	1.20	0.12	S	N	Y
6607357	7.700	67.390	0.85	0.30	2.72	0.45	0.55	0.22	5592	4.51	0.91	0.84	0.29	P1	N	N
6707835	22.248	84.881	2.25	0.23	1.97	0.07	< 0.43		5619	4.44	0.99	2.42	0.24	S	N	Y
6716545	13.910	75.855	0.84	0.30	3.29	0.68	< 0.85		6044	4.30	1.12	1.03	0.36	P1	N	N
6803202	21.061	76.591	1.58	0.16	2.64	0.08	< 0.36		5719	4.40	1.03	1.77	0.18	S	N	Y
6851425	11.120	72.744	2.29	0.23	1.71	0.12	< 0.62		5071	4.59	0.72	1.80	0.18	S	N	Y
6922710	23.127	78.663	1.01	0.11	2.57	0.47	< 0.78		5929	4.40	1.07	1.18	0.12	S	N	Y
7021534	9.066	68.147	1.42	0.50	0.81	0.05	< 0.55		5848	4.55	0.87	1.35	0.47	P1	Y	N
7033671	9.490	66.961	1.48	0.15	1.81	0.09	< 0.46		5679	4.29	1.26	2.03	0.20	S	N	Y
7211221	5.621	69.903	1.21	0.12	1.23	0.11	< 0.53		5634	4.44	0.93	1.23	0.12	S	N	Y
7219825	17.233	68.084	2.06	0.21	2.27	0.13	< 0.22		6089	4.36	1.13	2.54	0.25	S	N	Y
7345248	5.665	69.338	0.63	0.22	2.81	0.43	< 0.61		5656	4.27	1.19	0.81	0.29	P2		N
7419318	18.736	73.153	2.13	0.22	1.47	0.11	< 0.61		5187	4.54	0.81	1.89	0.19	S	N	Y
7466863	11.971	68.173	1.83	0.64	0.98	0.03	< 0.39		6035	4.12	1.45	2.89	1.01	P1	Y	N
7582689	11.921	70.475	0.75	0.26	5.84	1.24	< 0.95		6022	4.04	1.64	1.34	0.47	P1		N
7668663	6.498	69.012	1.44	0.15	1.85	0.34	0.79	0.09	5725	4.33	1.14	1.79	0.18	S	N	Y
7700622	35.585	86.426	2.83	0.28	1.99	0.10	< 0.46		4787	4.62	0.69	2.13	0.21	S	N	Y
7762723	9.887	72.539	0.73	0.26	2.21	0.39	< 0.65		5501	4.58	0.80	0.64	0.22	P1	Y	N
7810483	29.921	79.024	1.71	0.60	1.58	0.11	< 0.58		5893	4.52	0.91	1.70	0.59	P1	Y	N
7906739	7.015	69.903	0.90	0.31	2.43	0.32	< 0.68		5652	4.53	0.81	0.80	0.28	P1	Y	N
7906892	8.849	72.797	0.58	0.07	8.03	3.69	0.89	0.12	6095	4.35	1.14	0.72	0.08	S		N
7918652	11.456	69.316	0.79	0.28	2.45	0.47	< 0.81		5809	4.25	1.19	1.02	0.36	P1	N	N
8008067	15.771	70.584	2.22	0.22	3.32	0.31	0.67	0.08	5594	4.37	1.10	2.66	0.27	S	N	Y
8009496	38.476	83.567	1.88	0.66	1.72	0.14	< 0.70		5833	4.54	0.85	1.74	0.61	P1	Y	N
8073705	10.601	65.967	0.75	0.08	2.29	0.36	< 0.80		6086	4.36	1.12	0.91	0.10	S		N
8077137	15.090	78.772	0.78	0.08	2.47	0.65	< 0.85		6179	4.37	1.16	0.99	0.10	S	N	Y
8081187	37.323	85.828	1.67	0.59	3.58	0.39	< 0.62		6030	4.55	0.86	1.57	0.55	P1	Y	N
8087812	27.211	65.360	1.01	0.10	5.01	0.57	< 0.67		5985	4.17	1.41	1.55	0.16	S		N
8242434	44.964	77.565	2.58	0.26	2.58	0.15	0.31	0.22	4692	4.63	0.68	1.92	0.19	S	N	Y
8280511	10.435	67.826	1.37	0.14	1.70	0.12	< 0.58		5522	4.45	0.91	1.36	0.14	S		N
8323753	6.714	67.308	1.97	0.69	1.33	0.05	< 0.52		5817	4.23	1.26	2.71	0.95	P1	Y	N
8349582	11.523	64.959	2.14	0.22	2.41	0.21	0.61	0.09	5668	4.23	1.35	3.15	0.32	S	N	Y
8429668	5.007	67.696	0.73	0.26	1.66	0.47	0.62	0.29	5034	4.63	0.65	0.52	0.18	P2		N

Table 2 — Continued

kic	Light Curve Fit								Stellar Parameters							
	P (d)	t_0^a (d)	$\frac{R_P}{R_\star}$ (%)	$\sigma(\frac{R_P}{R_\star})$	τ (hrs)	$\sigma(\tau)$	b	$\sigma(b)$	T_{eff} (K)	$\log g$ (cgs)	R_\star (R_\odot)	R_P (R_E)	$\sigma(R_P)$	source ^b	FP	B12
8480285	29.667	92.687	2.43	0.25	4.72	0.42	0.48	0.14	5960	4.36	1.09	2.89	0.29	S	N	Y
8494617	22.923	66.640	1.06	0.11	3.72	0.36	< 0.50		5905	4.36	1.10	1.27	0.13	S	N	Y
8560804	31.976	66.803	0.99	0.35	4.96	0.48	< 0.65		5878	4.44	1.01	1.09	0.38	P1	N	N
8611832	22.597	73.431	1.08	0.11	3.15	0.17	< 0.58		5577	4.37	0.99	1.17	0.12	S	N	Y
8628758	14.374	71.204	1.99	0.20	4.76	0.68	< 0.91		5773	4.40	1.04	2.26	0.23	S	N	Y
8631504	14.820	66.409	1.20	0.12	2.02	0.25	< 0.63		4828	4.60	0.72	0.95	0.10	S	N	Y
8644365	19.917	72.354	1.07	0.11	3.37	0.55	< 0.85		6054	4.39	1.12	1.31	0.13	S	N	N
8804455	7.597	64.958	1.14	0.12	2.55	0.47	0.74	0.11	5715	4.38	1.07	1.33	0.14	S	N	Y
8805348	29.907	78.383	2.36	0.24	3.20	0.39	0.66	0.12	5739	4.34	1.05	2.69	0.27	S	N	Y
8822366	30.864	65.012	1.41	0.14	4.16	0.20	< 0.62		6089	4.35	1.14	1.76	0.18	S	N	Y
8827575	10.129	68.997	0.84	0.30	1.94	0.18	< 0.63		5284	4.45	0.89	0.82	0.29	P1	N	N
8866102	17.834	114.225	1.66	0.17	2.28	0.06	< 0.39		6178	4.37	1.15	2.08	0.21	S	N	Y
8962094	30.865	75.052	2.09	0.21	1.37	0.10	< 0.61		5739	4.34	1.04	2.38	0.24	S	N	Y
8972058	8.991	69.746	2.01	0.20	2.10	0.11	< 0.52		5979	4.38	1.09	2.39	0.24	S	N	Y
9006186	5.453	67.152	0.85	0.09	1.07	0.08	< 0.61		5404	4.53	0.87	0.81	0.08	S	N	Y
9086251	6.892	64.632	0.87	0.09	0.93	0.10	< 0.68		6044	4.22	1.45	1.38	0.14	S	N	Y
9139084	5.836	67.853	2.07	0.21	1.06	0.03	< 0.41		5411	4.53	0.85	1.92	0.19	S	N	Y
9226339	21.461	65.230	1.10	0.11	1.66	0.16	< 0.70		5807	4.28	1.25	1.50	0.15	S	N	N
9288237	7.491	68.329	0.52	0.18	3.01	0.73	< 0.84		5946	4.44	0.96	0.54	0.19	P1	N	N
9491832	49.565	103.693	1.14	0.12	6.43	1.49	0.74	0.25	5821	4.15	1.57	1.95	0.21	S	N	N
9549648	5.992	69.883	1.46	0.15	1.80	0.45	0.89	0.06	6165	4.38	1.14	1.82	0.19	S	N	Y
9704384	5.509	65.285	1.35	0.14	1.85	0.26	< 0.75		5448	4.50	0.91	1.35	0.14	S	N	Y
9716028	17.373	71.258	0.82	0.08	2.46	0.27	< 0.73		6119	4.37	1.15	1.03	0.11	S	N	N
9717943	6.110	69.903	0.72	0.08	1.80	0.87	0.79	0.20	5968	4.30	1.27	1.00	0.11	S	N	Y
9886361	7.031	67.464	0.88	0.09	3.04	0.21	< 0.42		6090	4.39	1.13	1.08	0.11	S	N	N
10055126	9.176	71.722	1.33	0.13	2.43	0.24	< 0.48		5905	4.36	1.10	1.59	0.16	S	N	Y
10130039	12.758	66.961	1.19	0.12	2.19	0.07	< 0.38		5828	4.42	1.01	1.31	0.13	S	N	Y
10136549	9.693	65.809	1.14	0.12	3.44	0.44	0.57	0.20	5684	4.13	1.59	1.98	0.20	S	N	Y
10212441	15.044	66.211	0.95	0.10	2.97	0.24	< 0.57		5939	4.35	1.09	1.13	0.11	S	N	Y
10593535	20.925	67.900	0.93	0.10	3.63	0.59	< 0.81		5822	4.28	1.25	1.27	0.13	S	N	N
10722485	7.849	67.907	0.87	0.09	2.61	0.51	< 0.87		5682	4.36	1.03	0.98	0.10	S	N	N
10917433	6.912	65.190	0.51	0.05	2.03	0.57	< 0.87		5680	4.33	1.12	0.62	0.06	S	N	N
11086270	31.720	75.821	1.90	0.19	3.35	0.40	0.70	0.14	5960	4.37	1.08	2.24	0.22	S	N	Y
11121752	7.630	70.087	1.00	0.10	1.56	0.23	< 0.75		6045	4.36	1.12	1.22	0.12	S	N	Y
11133306	41.746	101.657	1.90	0.19	2.32	0.15	0.27	0.22	5953	4.37	1.10	2.27	0.23	S	N	Y
11241912	14.427	71.857	0.95	0.10	2.40	0.18	< 0.59		5931	4.40	1.07	1.10	0.11	S	N	N
11250587	7.257	67.028	1.95	0.20	2.41	0.05	< 0.40		5853	4.18	1.49	3.18	0.32	S	N	Y
11253711	17.791	82.259	1.86	0.65	1.29	0.09	< 0.57		5816	4.48	0.95	1.92	0.67	P1	N	Y
11295426	5.399	69.065	1.89	0.19	2.78	0.20	0.80	0.05	5793	4.25	1.30	2.68	0.27	S	N	Y
11402995	10.061	71.959	2.00	0.20	2.42	0.19	< 0.62		5709	4.30	1.18	2.56	0.26	S	N	Y
11554435	9.434	73.119	5.69	0.57	1.29	0.02	< 0.43		5536	4.52	0.90	5.60	0.56	S	N	Y
11560897	35.968	71.667	1.46	0.15	1.47	0.11	< 0.45		5832	4.27	1.28	2.03	0.21	S	N	Y
11612280	9.406	70.699	0.80	0.08	3.02	0.54	< 0.77		5857	4.24	1.32	1.15	0.12	S	N	N
11771430	40.031	81.863	1.39	0.14	2.25	0.14	< 0.57		5850	4.23	1.38	2.10	0.21	S	N	Y
11774991	37.815	74.112	1.45	0.15	2.37	0.21	< 0.65		4710	4.62	0.70	1.10	0.11	S	N	Y
12254909	5.350	66.872	0.84	0.29	2.26	0.14	< 0.60		5987	4.46	0.96	0.88	0.31	P1	N	Y
12301181	6.147	67.867	1.04	0.11	1.45	0.07	< 0.40		4997	4.60	0.74	0.84	0.09	S	N	Y
12416661	8.053	67.968	0.63	0.22	2.93	0.37	< 0.59		6091	4.12	1.47	1.02	0.36	P1	N	N
12454461	7.467	69.392	0.84	0.09	1.83	0.20	< 0.71		6048	4.31	1.27	1.16	0.12	S	N	Y
12737015	24.669	69.616	1.05	0.11	4.94	0.49	< 0.69		6045	4.15	1.60	1.84	0.19	S	N	N

Note. — Orbital period, P ; time of transit center, t_0 ; planet-to-star radius ratio, R_P/R_\star ; the time for the planet to travel R_\star during transit, τ ; and transit impact parameter, b are all determined from the Mandel & Agol (2002) light curve fit. By default, stellar parameters R_\star , T_{eff} , and $\log g$ come from *SpecMatch*. If *SpecMatch* parameters do not exist, parameters are taken from the corrected KIC values, described in Section 4. The FP column lists whether a candidate was designated a false positive by the *Kepler* team (‘Y’-yes, ‘N’-no, ‘ ’-no designation). The B12 column lists whether a candidate was present in Batalha et al. (2012).

^a Time of transit center (BJD-2454900).

^b Source of stellar parameters: ‘S’-*SpecMatch*-derived parameters using Keck HIRES spectra, ‘P1’-photometrically-derived parameters from Batalha et al. (2012), ‘P2’-photometrically-derived parameters computed by the authors. See Section 4 for more details.

Table 3
New candidates identified with TERRA

kic	Light Curve Fit								Stellar Parameters					
	P (d)	t_0^a (d)	$\frac{R_P}{R_\star}$ (%)	$\sigma(\frac{R_P}{R_\star})$	τ (hrs)	$\sigma(\tau)$	b	$\sigma(b)$	T_{eff} (K)	$\log g$ (cgs)	R_\star (R_\odot)	R_P (R_E)	$\sigma(R_P)$	source ^b
3120904	42.915	72.866	1.15	0.12	2.96	0.36	< 0.66		6151	4.31	1.26	1.59	0.16	S

Table 3 — *Continued*

kic	Light Curve Fit								Stellar Parameters					source ^b
	P (d)	t_0^a (d)	$\frac{R_P}{R_\star}$ (%)	$\sigma(\frac{R_P}{R_\star})$	τ (hrs)	$\sigma(\tau)$	b	$\sigma(b)$	T_{eff} (K)	$\log g$ (cgs)	R_\star (R_\odot)	R_P (R_E)	$\sigma(R_P)$	
3545135	8.483	65.973	0.81	0.08	1.44	0.10	< 0.57		5794	4.40	1.02	0.90	0.09	S
3852655	11.629	65.817	0.85	0.30	1.49	0.23	< 0.79		5822	4.36	1.05	0.97	0.34	P1
4548011	6.284	66.198	0.60	0.06	2.23	0.64	< 0.49		5991	4.30	1.26	0.83	0.09	S
4770174	6.096	67.600	0.57	0.20	3.08	0.44	< 0.75		6013	4.44	1.01	0.63	0.22	P1
5096590	29.610	70.332	1.00	0.35	2.60	0.20	< 0.67		5623	4.63	0.73	0.79	0.28	P1
5308537	14.265	76.217	0.76	0.08	1.92	0.53	< 0.87		5831	4.29	1.26	1.05	0.11	S
5652893	14.010	65.959	1.15	0.13	1.98	0.49	< 0.85		5150	4.55	0.78	0.97	0.11	S
5702939	18.398	77.934	1.09	0.38	2.45	0.27	< 0.54		5634	4.47	0.87	1.04	0.36	P1
6197215	10.613	68.691	1.23	0.13	0.48	0.06	< 0.80		5933	4.39	1.09	1.46	0.15	S
6356692	11.392	74.555	0.66	0.23	2.74	0.31	< 0.73		5420	4.03	1.64	1.19	0.42	P1
6523351	6.067	69.128	0.72	0.25	1.02	0.20	< 0.89		5489	4.15	1.35	1.06	0.37	P1
6607357	7.700	67.390	0.85	0.30	2.72	0.45	0.55	0.22	5592	4.51	0.91	0.84	0.29	P1
6716545	13.910	75.855	0.84	0.30	3.29	0.68	< 0.85		6044	4.30	1.12	1.03	0.36	P1
7345248	5.665	69.338	0.63	0.22	2.81	0.43	< 0.61		5656	4.27	1.19	0.81	0.29	P2
7582689	11.921	70.475	0.75	0.26	5.84	1.24	< 0.95		6022	4.04	1.64	1.34	0.47	P1
7906892	8.849	72.797	0.58	0.07	8.03	3.69	0.89	0.12	6095	4.35	1.14	0.72	0.08	S
7918652	11.456	69.316	0.79	0.28	2.45	0.47	< 0.81		5809	4.25	1.19	1.02	0.36	P1
8073705	10.601	65.967	0.75	0.08	2.29	0.36	< 0.80		6086	4.36	1.12	0.91	0.10	S
8087812	27.211	65.360	1.01	0.10	5.01	0.57	< 0.67		5985	4.17	1.41	1.55	0.16	S
8280511	10.435	67.826	1.37	0.14	1.70	0.12	< 0.58		5522	4.45	0.91	1.36	0.14	S
8429668	5.007	67.696	0.73	0.26	1.66	0.47	0.62	0.29	5034	4.63	0.65	0.52	0.18	P2
8560804	31.976	66.803	0.99	0.35	4.96	0.48	< 0.65		5878	4.44	1.01	1.09	0.38	P1
8644365	19.917	72.354	1.07	0.11	3.37	0.55	< 0.85		6054	4.39	1.12	1.31	0.13	S
8827575	10.129	68.997	0.84	0.30	1.94	0.18	< 0.63		5284	4.45	0.89	0.82	0.29	P1
9226339	21.461	65.230	1.10	0.11	1.66	0.16	< 0.70		5807	4.28	1.25	1.50	0.15	S
9288237	7.491	68.329	0.52	0.18	3.01	0.73	< 0.84		5946	4.44	0.96	0.54	0.19	P1
9491832	49.565	103.693	1.14	0.12	6.43	1.49	0.74	0.25	5821	4.15	1.57	1.95	0.21	S
9716028	17.373	71.258	0.82	0.08	2.46	0.27	< 0.73		6119	4.37	1.15	1.03	0.11	S
9886361	7.031	67.464	0.88	0.09	3.04	0.21	< 0.42		6090	4.39	1.13	1.08	0.11	S
10593535	20.925	67.900	0.93	0.10	3.63	0.59	< 0.81		5822	4.28	1.25	1.27	0.13	S
10722485	7.849	67.907	0.87	0.09	2.61	0.51	< 0.87		5682	4.36	1.03	0.98	0.10	S
10917433	6.912	65.190	0.51	0.05	2.03	0.57	< 0.87		5680	4.33	1.12	0.62	0.06	S
11241912	14.427	71.857	0.95	0.10	2.40	0.18	< 0.59		5931	4.40	1.07	1.10	0.11	S
11612280	9.406	70.699	0.80	0.08	3.02	0.54	< 0.77		5857	4.24	1.32	1.15	0.12	S
12416661	8.053	67.968	0.63	0.22	2.93	0.37	< 0.59		6091	4.12	1.47	1.02	0.36	P1
12737015	24.669	69.616	1.05	0.11	4.94	0.49	< 0.69		6045	4.15	1.60	1.84	0.19	S

Note. — The 37 TERRA candidates not in Batalha et al. (2012) and not listed as false positives by the *Kepler* team. The column definitions are the same as in Table 2

Table 4
Union of Batalha et al. (2012) and TERRA planet candidate catalogs

kic	Batalha			TERRA	
	KOI	P	R_P	P	R_P
2142522				13.32	1.11
2307415	2053.01	13.12	1.65	13.12	1.57
2441495	166.01	12.49	2.70	12.49	1.99
2444412	103.01	14.91	2.97	14.91	3.41
2571238	84.01	9.29	2.53	9.29	2.82
2853446	1118.01	7.37	2.52	7.37	1.65
3098810	1878.01	40.81	3.24	40.81	2.75
3120904				42.91	1.59
3342794	2278.01	14.17	1.99	14.17	1.68
3442055	1218.01	29.62	2.22	29.62	1.72
3531558	118.01	24.99	1.41	24.99	1.85
3545135				8.48	0.90
3835670	149.01	14.56	5.50	14.56	4.93
3839488	1216.01	11.13	1.57	11.13	1.64
3852655				11.63	0.97
3942670	392.02	12.61	1.33		
3942670	392.01	33.42	2.27	33.42	2.13
4043190	1220.01	6.40	1.95	6.40	2.86
4049901	2295.01	16.29	0.63	16.29	0.60
4548011				6.28	0.83
4644604	628.01	14.49	1.87	14.49	2.32
4770174				6.10	0.63
4827723	632.01	7.24	1.46	7.24	1.54

Table 4 — *Continued*

kic	Batalha			TERRA	
	KOI	P	R_P	P	R_P
4914423	108.01	15.97	2.94	15.97	3.05
4914566				22.24	1.18
5009743	1609.01	41.70	2.34	41.70	2.32
5042210	2462.01	12.15	1.37	12.15	1.16
5094751	123.01	6.48	2.64	6.48	2.02
5094751	123.02	21.22	2.71		
5096590				29.61	0.79
5121511	640.01	31.00	2.43	31.00	2.15
5308537				14.27	1.05
5446285	142.01	10.92	4.08		
5561278	1621.01	20.31	2.48	20.31	1.45
5613330	649.01	23.45	2.31	23.45	2.70
5652893				14.01	0.97
5702939				18.40	1.04
5735762	148.02	9.67	3.14	9.67	2.51
5735762	148.03	42.90	2.35		
5866724	85.01	5.86	2.35	5.86	2.37
5866724	85.03	8.13	1.41		
5959719	2498.01	6.74	0.78	6.74	0.83
6061773	2001.01	8.28	2.39		
6071903	306.01	24.31	2.28	24.31	2.06
6197215				10.61	1.46
6289257	307.02	5.21	1.15		
6289257	307.01	19.67	1.80	19.67	1.61
6291837	308.01	35.60	3.15	35.60	3.05
6356692				11.39	1.19
6365156	662.01	10.21	2.05	10.21	2.13
6442340	664.02	7.78	1.19		
6442340	664.01	13.14	1.83	13.14	1.63
6442340	664.03	23.44	1.19		
6521045	41.02	6.89	1.23		
6521045	41.01	12.82	2.08	12.82	1.85
6521045	41.03	35.33	1.40		
6523351				6.07	1.06
6605493	2559.01	9.31	0.99	9.31	1.20
6607357				7.70	0.84
6678383	111.01	11.43	2.14		
6678383	111.02	23.67	2.05		
6707835	666.01	22.25	2.56	22.25	2.42
6716545				13.91	1.03
6803202	177.01	21.06	1.84	21.06	1.77
6850504	70.04	6.10	0.91		
6850504	70.01	10.85	3.09		
6850504	70.05	19.58	1.02		
6851425	163.01	11.12	2.27	11.12	1.80
6922710	2087.01	23.13	1.54	23.13	1.18
7021534				9.07	1.35
7033671	670.01	9.49	1.92	9.49	2.03
7211221	1379.01	5.62	1.06	5.62	1.23
7219825	238.01	17.23	2.40	17.23	2.54
7219825	238.02	26.69	1.35		
7345248				5.66	0.81
7419318	313.02	8.44	1.61		
7419318	313.01	18.74	2.20	18.74	1.89
7466863				11.97	2.89
7582689				11.92	1.34
7668663	1898.01	6.50	1.50	6.50	1.79
7700622	315.01	35.59	2.14	35.59	2.13
7762723				9.89	0.64
7810483				29.92	1.70
7906739				7.01	0.80
7906892				8.85	0.72
7918652				11.46	1.02
8008067	316.01	15.77	2.72	15.77	2.66
8009496				38.48	1.74
8073705				10.60	0.91
8077137	274.01	15.09	1.12	15.09	0.99
8077137	274.02	22.80	1.13		
8081187				37.32	1.57
8087812				27.21	1.55
8242434	1726.01	44.96	5.25	44.96	1.92
8280511	1151.01	5.22	0.84		
8280511	1151.02	7.41	0.97		
8280511				10.44	1.36
8323753				6.71	2.71

Table 4 — *Continued*

kic	Batalha			TERRA	
	KOI	P	R_P	P	R_P
8349582	122.01	11.52	2.78	11.52	3.15
8429668				5.01	0.52
8480285	691.02	16.23	1.25		
8480285	691.01	29.67	2.92	29.67	2.89
8494617	2389.01	22.92	1.45	22.92	1.27
8554498	5.02	7.05	0.66		
8560804				31.98	1.09
8611832	2414.01	22.60	1.03	22.60	1.17
8611832	2414.02	45.35	1.17		
8628758	1279.02	9.65	0.74		
8628758	1279.01	14.37	1.31	14.37	2.26
8631504	2503.01	14.82	2.41	14.82	0.95
8644365				19.92	1.31
8804455	2159.01	7.60	1.01	7.60	1.33
8805348	695.01	29.91	2.51	29.91	2.69
8822366	1282.01	30.86	3.00	30.86	1.76
8827575				10.13	0.82
8866102	42.01	17.83	2.71	17.83	2.08
8962094	700.02	9.36	1.29		
8962094	700.03	14.67	1.29		
8962094	700.01	30.86	2.28	30.87	2.38
8972058	159.01	8.99	2.70	8.99	2.39
9006186	2169.01	5.45	1.02	5.45	0.81
9086251	2367.01	6.89	1.17	6.89	1.38
9139084	323.01	5.84	2.17	5.84	1.92
9226339				21.46	1.50
9288237				7.49	0.54
9471974	119.01	49.18	3.76		
9491832				49.57	1.95
9549648	1886.01	5.99	2.45	5.99	1.82
9704384	1913.01	5.51	1.40	5.51	1.35
9716028				17.37	1.03
9717943	2273.01	6.11	1.02	6.11	1.00
9886361				7.03	1.08
10055126	1608.01	9.18	1.81	9.18	1.59
10055126	1608.02	19.74	1.58		
10130039	1909.02	5.47	1.15		
10130039	1909.01	12.76	1.52	12.76	1.31
10130039	1909.03	25.10	1.63		
10136549	1929.01	9.69	2.00	9.69	1.98
10212441	2342.01	15.04	1.22	15.04	1.13
10593535				20.92	1.27
10722485				7.85	0.98
10917433				6.91	0.62
11086270	124.01	12.69	3.00		
11086270	124.02	31.72	3.58	31.72	2.24
11121752	2333.02	7.63	1.63	7.63	1.22
11133306	276.01	41.75	2.49	41.75	2.27
11241912				14.43	1.10
11250587	107.01	7.26	3.09	7.26	3.18
11253711	1972.01	17.79	1.93	17.79	1.92
11295426	246.01	5.40	2.53	5.40	2.68
11402995	173.01	10.06	2.48	10.06	2.56
11554435	63.01	9.43	6.30	9.43	5.60
11560897	2365.01	35.97	1.59	35.97	2.03
11612280				9.41	1.15
11771430	2582.01	40.03	1.98	40.03	2.10
11774991	2173.01	37.82	1.24	37.82	1.10
11818872	2581.01	12.74	0.90		
12254909	2372.01	5.35	1.11	5.35	0.88
12301181	2059.01	6.15	0.59	6.15	0.84
12416661				8.05	1.02
12454461	2463.01	7.47	1.07	7.47	1.16
12737015				24.67	1.84

Note. — All Batalha et al. (2012) candidates with $P = 5$ –50 days belonging to stars in the Best12k sample are included. Candidates are considered equal if they belong to the same star and the periods in each catalog agree to better than 0.01 days. Eighty-two candidates appear in both catalogs, 33 appear in Batalha et al. (2012) only, and 47 appear in this work only (although 10 were listed as false positives by the *Kepler* team). Differences in R_P between the two catalogs stem from different values of R_* . Most TERRA planet candidates have SpecMatch-derived stellar parameters which are more accurate than Batalha et al. (2012) parameters, which were derived from KIC broadband photometry.

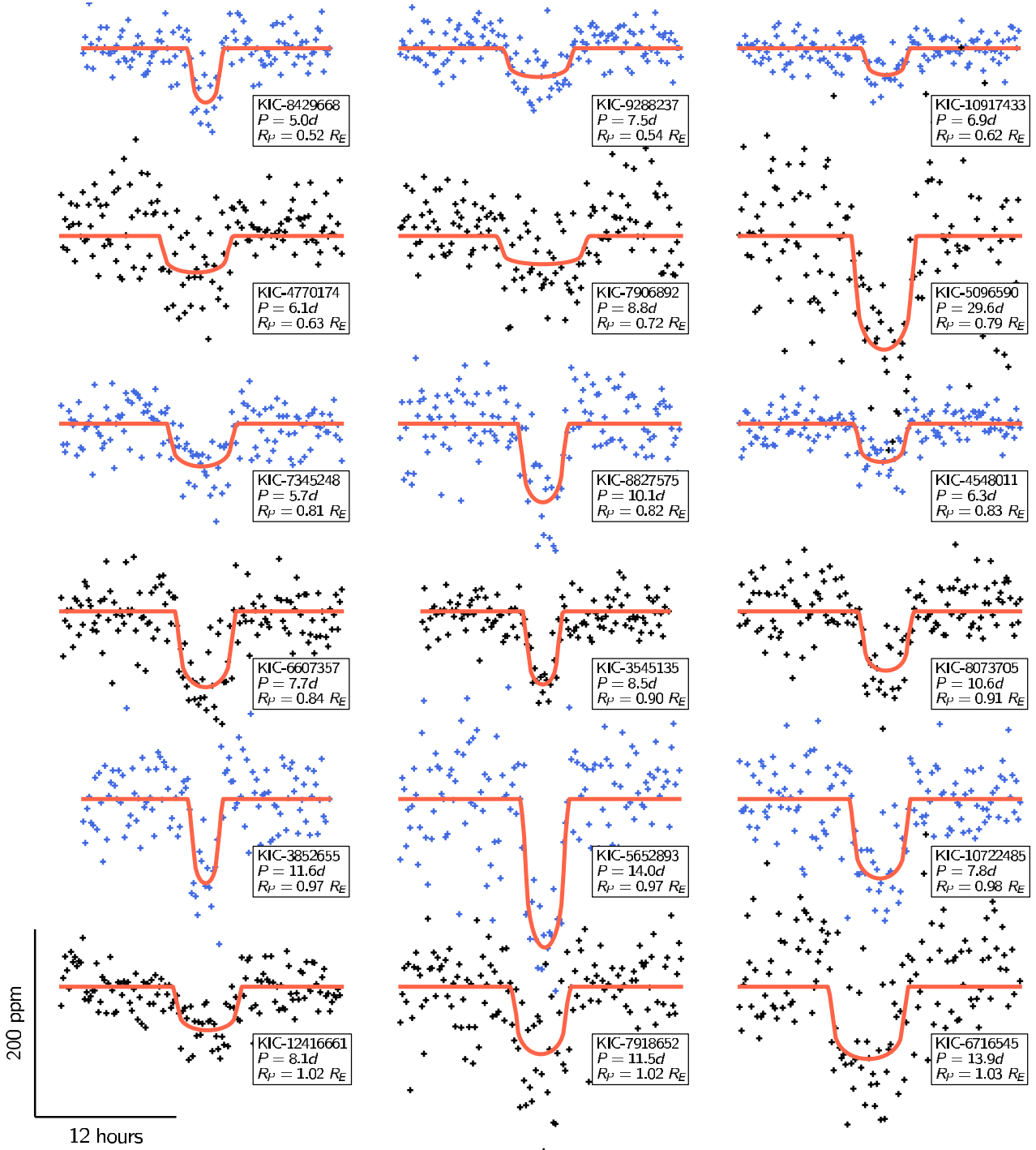


Figure 17. Phase-folded photometry for 18 of the 37 TERRA planet candidates, not in Batalha et al. (2012), ordered according to size. For clarity, we show median photometric measurements in 10 min bins. The red lines are the best-fitting Mandel & Agol (2002) model.

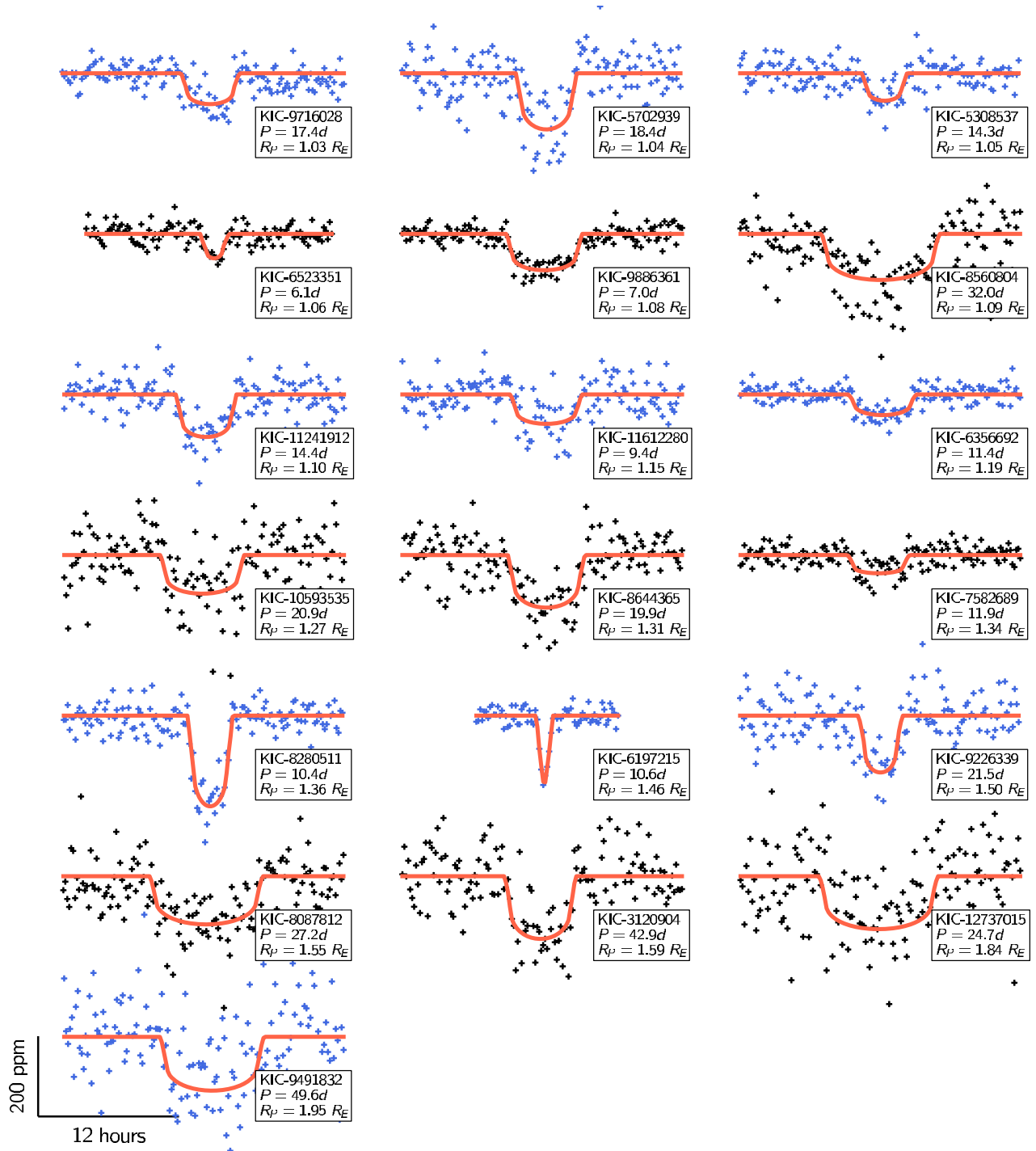


Figure 18. Same as Figure 17, but showing the remaining 19 of the 37 TERRA planet candidates, not in Batalha et al. (2012).