

ACCRETION ONTO PLANETARY MASS COMPANIONS OF LOW-MASS YOUNG STARS

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ABSTRACT

Measurements of accretion rates onto planetary mass objects may distinguish between different planet formation mechanisms, which predict different accretion histories. In this Letter, we use *Hubble Space Telescope* (*HST*)/WFC3 UVIS optical photometry to measure accretion rates onto three accreting objects, GSC 06214–00210 b, GQ Lup b, and DH Tau b, that are at the planet/brown dwarf boundary and are companions to solar mass stars. The excess optical emission in the excess accretion continuum yields mass accretion rates of 10^{-9} – $10^{-11} M_{\odot} \text{ yr}^{-1}$ for these three objects. Their accretion rates are an order of magnitude higher than expected from the correlation between mass and accretion rates measured from the UV excess, which is applicable if these wide planetary mass companions formed by protostellar core fragmentation. The high accretion rates and large separation from the central star demonstrate the presence of massive disks around these objects. Models for the formation and evolution of wide planetary mass companions should account for their large accretion rates. High ratios of $H\alpha$ luminosity over accretion luminosity for objects with low accretion rates suggest that searches for $H\alpha$ emission may be an efficient way to find accreting planets.

Key words: accretion, accretion disks – brown dwarfs – protoplanetary disks – stars: low-mass – stars: pre-main sequence

1. INTRODUCTION

In recent years, many young objects with planetary masses have been directly imaged at distances of tens to a few hundred AU from the central star (e.g., Neuhäuser et al. 2005; Itoh et al. 2005; Kraus et al. 2008; Marois et al. 2008; Kraus & Ireland 2012; Quanz et al. 2013). The formation of these planetary mass companions may occur by planetary core formation followed by vigorous gas accretion, by gravitational instabilities within the disk, or by uneven fragmentation of the collapsing protostellar core. These formation mechanisms each predict different accretion histories, which lead to different contraction timescales and luminosities at ages <100 Myr (Spiegel & Burrows 2012; Mordasini 2013). Since younger objects are brighter and therefore more easily detected, the uncertain formation and accretion history leads to significant uncertainties in derived masses of directly imaged exoplanets.

If very low mass companions have similar disk fractions as solar mass stars and have accretion rates that are expected for their mass, then their formation may be consistent with the low-mass tail of the initial mass function resulting from protostellar core fragmentation. A similar approach established that free floating brown dwarfs likely form by the collapse of a protostellar core, similar to solar mass stars (e.g., White & Basri 2003; Muzerolle et al. 2003; Mohanty et al. 2005; Joergens et al. 2013). On the other hand, planetary mass objects that form by core accretion or by disk gravitational instabilities may have different accretion and disk properties. Core accretion models for giant planet formation predict that for a few Myr, the accretion rates of planetary companions should be very large and may even dominate the flux for a few Myr (Spiegel & Burrows 2012; Mordasini 2013).

Previously, accretion has been detected onto two very low mass companions of solar mass stars based on emission in the Paschen- β line (e.g., Seifahrt et al. 2007; Bowler et al. 2011; Bonnefoy et al. 2013). The line luminosities may then be converted to accretion rate using established correlations (Natta et al. 2006); however the uncertainties and scatter in the correlations introduce significant uncertainties in the resulting accretion rates, especially beyond the mass accretion rate regime over which the correlation was calculated. In this Letter, we use broadband optical *Hubble Space Telescope* (*HST*)/WFC3 photometry to directly measure the accretion rate of very low mass companions to the pre-main sequence stars GSC 06214–00210, GQ Lup, and DH Tau. The accretion luminosity is measured directly from the excess line and continuum emission, following the approach of Herczeg & Hillenbrand (2008). The use of photometry to obtain accretion rates is similar to the approach of Hartmann et al. (1998) and White & Ghez (2001).

2. OBSERVATIONS

We used *HST*/WFC3 UVIS2 to obtain optical photometry of GQ Lup b on 2012 February 25, GSC 06214–00210 b on 2012 February 22, and DH Tau b on 2012 January 22 in *HST* program GO 12507 (PI: A. Kraus). The high spatial resolution of *HST* is typically needed to resolve the components at optical wavelengths. The reduced and flatfielded images were downloaded from the MAST archive for analysis. The observation log and extracted fluxes are listed in Table 1.

2.1. Sample Selection and Properties

The total data set consists of optical imaging of the 12 planetary mass companions to young solar mass stars that had been

Table 1
Observations

WFC3 Filter	t_{exp} (s)	Mag	Flux ^a (erg cm ⁻² s ⁻¹)	Error
GSC 6214–210 b				
F336W	2400	21.9	5.55E-18	1.4%
F390W	480	22.1	8.16E-18	1.6%
F475W	320	21.9	4.23E-18	5.5%
F555W	240	22.3	4.62E-18	6.2%
F625W	140	20.8	1.13E-17	3.2%
(F625W) ^b		21.3	7.33E-18	3.2%
F656N	2380	15.7	7.08E-16	0.35%
F673N	800	21.5	4.66E-18	7.2%
F775W	80	20.2	1.05E-17	2.9%
F850LP	50	18.4	3.32E-17	2.0%
		$\log L_{\text{slab}}/L_{\odot}$	–4.65	
		$\log L_{H-\alpha}/L_{\odot}$	–5.03	
		EW _{Hα}	1600 Å	
GQ Lup b				
F336W	400	19.2	6.77E-17	1.6%
F390W	50	20.1	5.42E-17	2.0%
F475W	40	20.0	2.61E-17	5.2%
F555W	30	20.2	3.19E-17	3.5%
F625W	20	18.9	6.88E-17	8.0%
(F625W)		19.2	5.03E-17	8.9%
F656N	250	15.9	5.92E-16	3.9%
F673N	160	19.0	4.84E-17	15.3%
F775W	40	17.8	9.60E-17	2.9%
F850LP	20	16.2	2.75E-16	1.8%
		$\log L_{\text{slab}}/L_{\odot}$	–2.91	
		$\log L_{H-\alpha}/L_{\odot}$	–4.69	
		EW _{Hα}	180 Å	
DH Tau b				
F336W	1400	24.2	6.90E-19	30%
F390W	360	24.9	6.62E-19	20%
F475W	280	24.3	4.85E-19	9.0%
F555W	160	24.5	5.98E-19	7.2%
F625W	100	23.0	1.51E-18	5.7%
(F625W)		23.1	1.35E-18	6.6%
F656N	1200	19.0	3.48E-17	3.4%
F673N	500	23.0	1.23E-18	21%
F775W	80	20.2	1.07E-17	2.7%
F850LP	40	18.0	5.16E-17	1.8%
		$\log L_{\text{slab}}/L_{\odot}$	–5.40	
		$\log L_{H-\alpha}/L_{\odot}$	–6.19	
		EW _{Hα}	450 Å	

Notes.

^a Observed fluxes not corrected for extinction. Listed luminosities are corrected for extinction.

^b F625W fluxes in parentheses are fluxes with H α emission subtracted.

identified as of 2011 February, when the proposal was submitted. In this Letter, we analyze three objects, GQ Lup b (Neuhäuser et al. 2005), GSC 06214–00210 b (Kraus et al. 2008), and DH Tau b (Itoh et al. 2005), that show excess emission in H α and in the *U*-band. A subsequent paper will describe the full sample. Emission in Pa β , an accretion diagnostic, has previously been detected from GSC 06214–00210 b and DH Tau b (Bowler et al. 2011; Bonnefoy et al. 2013). Pa β emission was detected in one of three near-IR spectra of GQ Lup b (Seifahrt et al. 2007; McElwain et al. 2007; Lavigne et al. 2009).

The properties of the three planetary mass companions are listed in Table 2. The extinctions of the companion are assumed to be equal to that of the primary. The extinctions to the accreting stars GQ Lup and DH Tau are calculated from optical spectra by Herczeg & Hillenbrand (2014). The extinction to the non-accreting star GSC 06214–210 is calculated here from the *R* – *J* color. Extinction curves applied here assume $R_V = 3.1$ (Fitzpatrick & Massa 1990). The distances and ages of GSC 06214–00210, GQ Lup, and DH Tau are assumed to be that of their parent clouds: Upper Sco OB Association at 145 pc (de Zeeuw et al. 1999) and 5–12 Myr (Preibisch et al. 2002; Pecaut et al. 2012), Lupus 1 at 155 pc (Lombardi et al. 2008) and 3 Myr (Alcala et al. 2014), and Taurus at 145 pc (Torres et al. 2009) and 1–2 Myr (e.g., Luhman 2004), respectively. The listed separations and position angles (P.A.s) are calculated from the median of our measurements for each filter with the plate scale of 0''.040 pixel⁻¹.

2.2. Source Extraction

The images consist of bright emission from the primary star and faint emission from the companion. Isolated stars in the field are used to synthesize a template point-spread function (PSF) in each image. For the GSC 06214–00210 images obtained with the F656N, F673N, and F850LP filters, the template PSF is first scaled to the primary and subtracted from the image, leaving only the faint companion. For other images of GSC 06214–00210 A and all the images of GQ Lup A and DH Tau A, the primary stars are heavily saturated and poorly fit with the template. The saturated spot is assumed to be symmetric about the *x*- and *y*-axis and rotate the stellar image to subtract out the primary. Figure 1 demonstrates our PSF subtraction for the F656N observations of all three objects.

After the primary flux is subtracted, the flux in the secondary component measured by fitting the spot with the template PSF. The template brightness was measured from aperture photometry and converted to a flux based on the aperture correction and flux conversion listed in the WFC3 manual.⁸ The extraction apertures have radii as large as 30 pixels (1''.2) for images with bright secondaries well separated from the primary, to as little as 3 pixels (0''.12) for faint images. The fluxes of the PSFs are then corrected by applying the flux-aperture radius curve from the WFC3 manual.

The photometry of the companion is then calculated by fitting the PSF template to the image of the companion. GQ Lup b and GSC 06214–0210 b are located close to the diffraction spikes, which are avoided in the two-dimensional fits to the image. The uncertainties are calculated from statistical error and fluctuation of the background in an annular region around the primary star.

The final photometric results and 1 σ uncertainties are listed in Table 1. The H α fluxes are calculated by convolving a flux above the model fit with the filter transmission curve to obtain the measured flux. The approximate equivalent widths are calculated by dividing the measured H α flux by the average flux in the F625W and F673N filters.

A faint object is detected in only the F656N image of DH Tau, with flux of 2.4×10^{-17} erg cm⁻² s⁻¹ Å⁻¹, a separation of 1''.53 and a P.A. of 46°.37 relative to the primary star. This object has upper limits of 2.1×10^{-19} and 3.7×10^{-19} erg cm⁻² s⁻¹ Å⁻¹ in the F775W and F850LP filters. A faint object is also found at a

⁸ http://www.stsci.edu/hst/wfc3/documents/handbooks/currentIHB/c06_uvis07.html#391868

Table 2
Sample Properties

Object	Sep. ($''$)	P.A. (deg)	d (pc)	A_V	T_{eff} (K)	$\log L_{\text{phot}}$	R (R_J)	M (M_J)	$\log L_{\text{acc}}$	$\log \dot{M}$	Ref.
GSC 06214 –00210 b	2.19 ± 0.01	175.2 ± 0.2	145 ± 15	0.2	2700 ± 200 2200	-3.1 ± 0.1 –3.16	1.5 ± 0.5 1.8 ± 0.5	14 ± 2 15 ± 3	–4.7 –4.6	–11.0 –10.8	1 ...
GQ Lup b	0.713 ± 0.006	-83.6 ± 0.7	155 ± 15	1.5	2050 ± 350 2400	-2.25 ± 0.24 –2.24	6.5 ± 2 4.6 ± 1.4	24 ± 12 31 ± 10	–2.9 –2.9	–9.3 –9.3	2, 3 ...
DH Tau b	2.31 ± 0.02	138.5 ± 0.1	145 ± 15	0.7	2350 ± 150 2200	-2.70 ± 0.11 –2.82	2.6 ± 0.7 2.7 ± 0.8	11_{-3}^{+10} 11 ± 3	–5.4 –5.3	–11.5 –11.3	2, 4 ...

Notes. For each object, the first line refers to literature T_{eff} , L_{phot} , R , and M values. The second line is measured here. The $L_{\text{acc}} = L_{\text{slab}} + L_{\text{H}\alpha}$ and \dot{M} are measured for the photospheric parameter corresponding to each line. For a description of the uncertainty in \dot{M} , see Section 3.

References. (1) Bowler et al. 2011; (2) Patience et al. 2012; (3) Seifahrt et al. 2007; (4) Itoh et al. 2005.

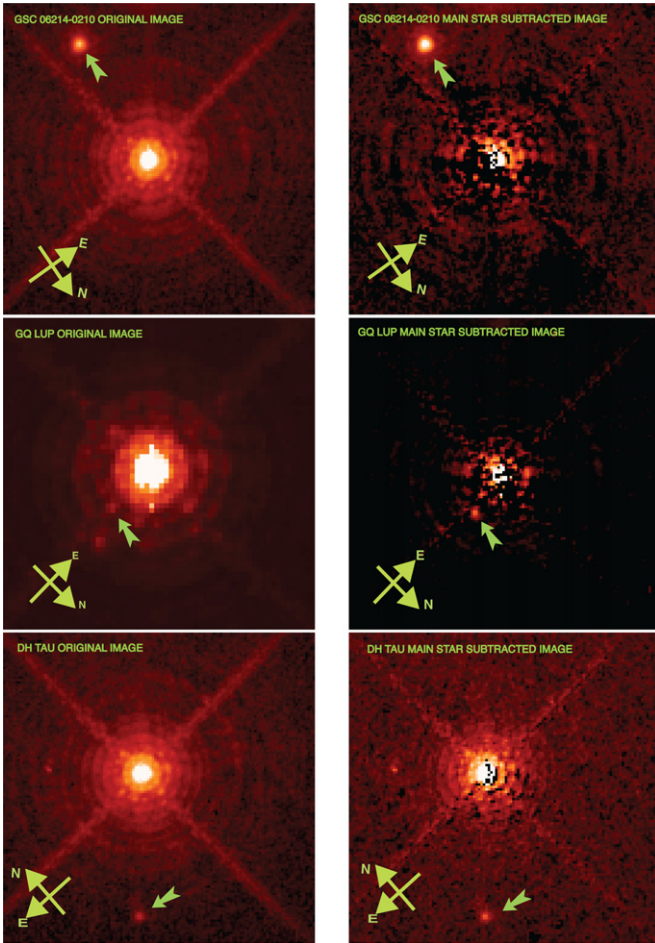


Figure 1. Subtracting the primary in the GSC 06214–00210 F656N image (top), the GQ Lup image (middle) F656N, and DH Tau (bottom) F656N. For GSC 06214–00210 and DH Tau, the primary star is subtracted by a scaled PSF. The secondary is then fit with a model PSF for GQ Lup. The primary star from the original image (left) is rotated by 180° and subtracting, leaving a clearer image of the secondary. The bright point on the left of the primary star of DH Tau only appears on the image of F656N.

similar relative position to DI Tau in the same image, indicating that both detections are likely artifacts.

3. RESULTS

Figure 2 shows our optical photometry combined with *JHK* photometry of GQ Lup and DH Tau from Patience et al. (2012) and GSC 06214–00210 from Bowler et al. (2011). The sources

are all brightest at red wavelengths, as expected for very low mass sources. However, the optical fluxes are relatively flat with wavelength, and higher in the F330W and F390W filters than at longer wavelengths. The emission in the narrowband F656N filter is much brighter than the nearby spectral regions because of $\text{H}\alpha$ emission, which also contributes 10%, 25%, and 77% of the photons in the broadband F625W filter for GQ Lup, FW Tau, and GSC 06214–0210, respectively.

The spectrophotometry is explained as the combination of a photosphere from a very low mass object and an accretion spectrum. Accretion is characterized by excess optical emission in the H Paschen continuum, a stronger excess in the H Balmer continuum shortward of 3700 \AA , and strong emission in H lines (see, e.g., Calvet & Gullbring 1998; Herczeg et al. 2009). The brightness of emission in the F330W filter is interpreted as the excess Balmer continuum emission. Emission in high Balmer lines and Ca II H and K likely explains the bright emission in the F390W filter, relative to longer wavelengths. Other optical lines may be strong enough to contaminate emission in other filters (e.g., Bowler et al. 2014), but are assumed to be negligible. In particular, the F673N filter is consistent with expectations of the photospheric flux, indicating at most a minimal contribution from the [S II] 6725 \AA doublet.

The photometry is modeled as the sum of the object photosphere and an accretion spectrum. Synthetic photometry for the photosphere is obtained from the BT-Settl grid with CIFIST opacities (Allard et al. 2012). The accretion spectrum is assumed to be flat, in $\text{erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$, with wavelength for all filters longer than 4000 \AA , following Herczeg & Hillenbrand (2014). The bolometric correction to convert the excess accretion continuum is calculated from the pure hydrogen slab models developed by Valenti et al. (1993). These slab models are less physically realistic than shock models (Calvet & Gullbring 1998; Ingleby et al. 2013), but both approaches provide similar bolometric corrections to obtain the final accretion luminosity.

These model spectra are then combined to reproduce the combined optical/near-IR photometry using χ^2 fits with two different methods. In the first set of fits, the photospheric luminosities are fixed using *JHK* photometry and effective temperatures are fixed to the values obtained from near-IR spectra (Table 2). The only free parameter is the excess luminosity of the slab. In the second set of fits, the temperature, photospheric luminosity, and excess accretion luminosity are all treated as free parameters. The best fit effective temperatures and luminosities are listed in Table 2. Radii and the photospheric luminosities of the three object are calculated by scaling the BT-Settl synthetic spectra to the observed spectral energy distribution (SED). The

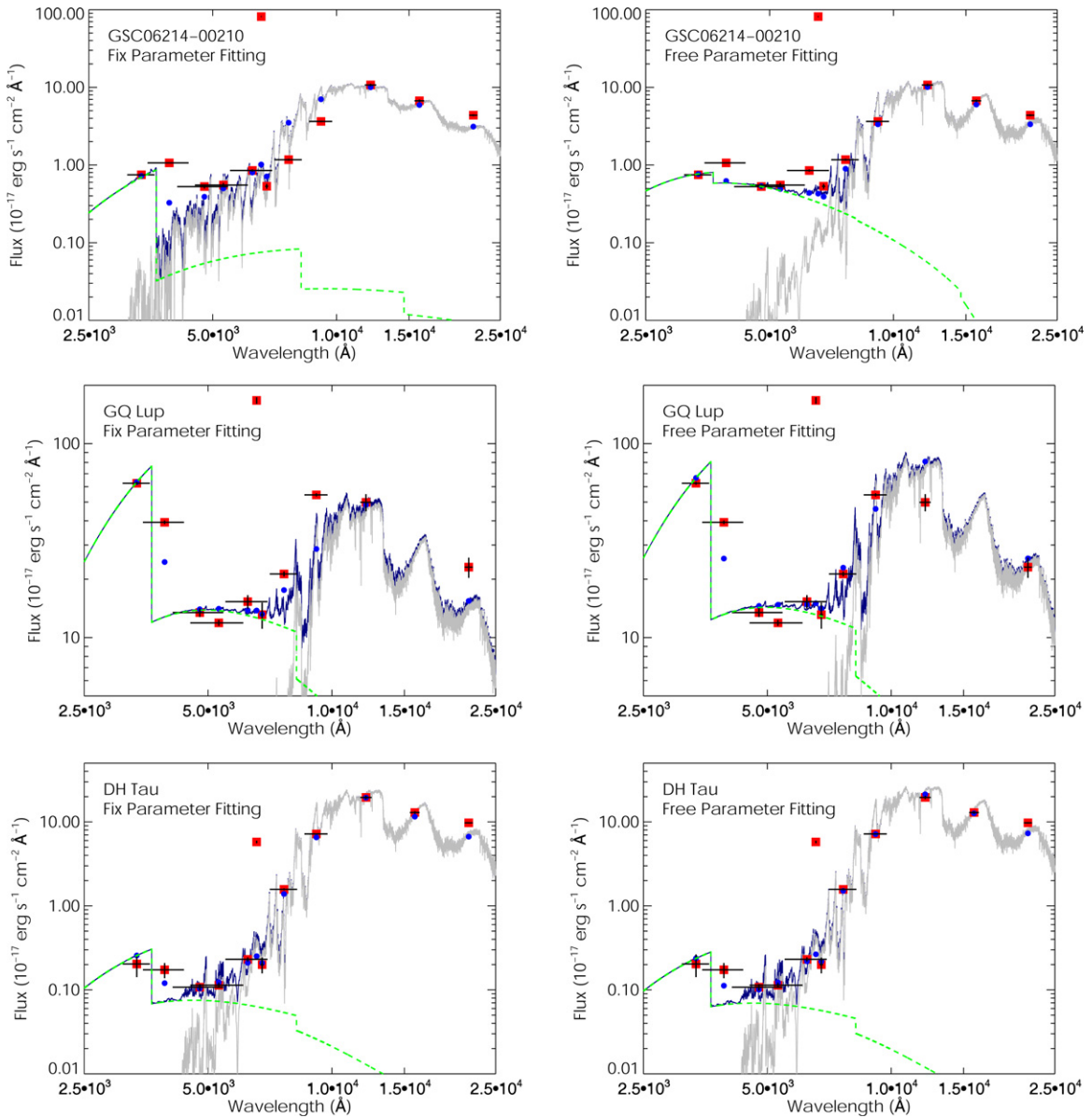


Figure 2. Optical and NIR SED with BT-Settl grid and slab emission of GSC 06214–00210 (top), GQ Lup b (middle), and DH Tau b (bottom). The left figures are the results for fixed parameter fitting. The right figures are for free parameter fitting. Red square points are observed photometry results. Green dashed lines are emission from a pure hydrogen layer, gray lines are the fluxes from BT-Settl model spectra, dark blue lines are the total flux of photosphere and accretion model. The integrate flux of the models are plotted in blue points. The SEDs present clear optical excess emissions. H α emissions in F625W filter have been subtracted.

uncertainty in radius is derived from the uncertainties in extinction and distance. The range in acceptable mass is set by fitting with pre-main sequence evolution model following Figure 5 in Bowler et al. (2011).

The spectral fits to DH Tau b and GSC 06214–0210 b are well matched to most of the observed photometric points. Possible K band excesses of 0.3–0.4 mag could be explained by emission from the warm disk, similar to the L -band excess around GSC 06214–00210 (Ireland et al. 2011). Alternatively, young substellar and planetary mass objects tend to be systematically redder than model spectra because of the gravity dependence of dust grain sizes and the height of clouds (e.g., Marley et al. 2012).

The J -band and F850LP photometry of GQ Lup b have similar fluxes, which is not expected from very cool objects and is poorly fit with the model spectra. In fits with luminosity and temperature as free parameters, no K -band excess is detected by the J -band flux is severely overestimated by the models. This

discrepancy is unlikely to be explained by excess line emission in the F850LP filter, because in that case more lines would have been detected in the near-IR spectra. Large photometric variability is often detected in accreting T Tauri stars (e.g., Herbst et al. 1994) and may affect the non-simultaneous comparison between the optical and near-IR photometry. The small separation between GQ Lup A and its companion makes the primary star subtraction challenging and could introduce photometric errors.

The effective temperatures measured here differ by a few hundred K from the literature values. The fits of synthetic spectra to near-IR spectra (e.g., Patience et al. 2012) should yield more accurate temperatures than our fits to optical+near-IR photometry. However, both temperatures are listed in case the depth of molecular bands is affected by low gravity in ways that are not yet accounted for in the models, in which case broadband SED fits would yield better temperatures. Spectral features may be weakened by any veiling from the disk in the K band and,

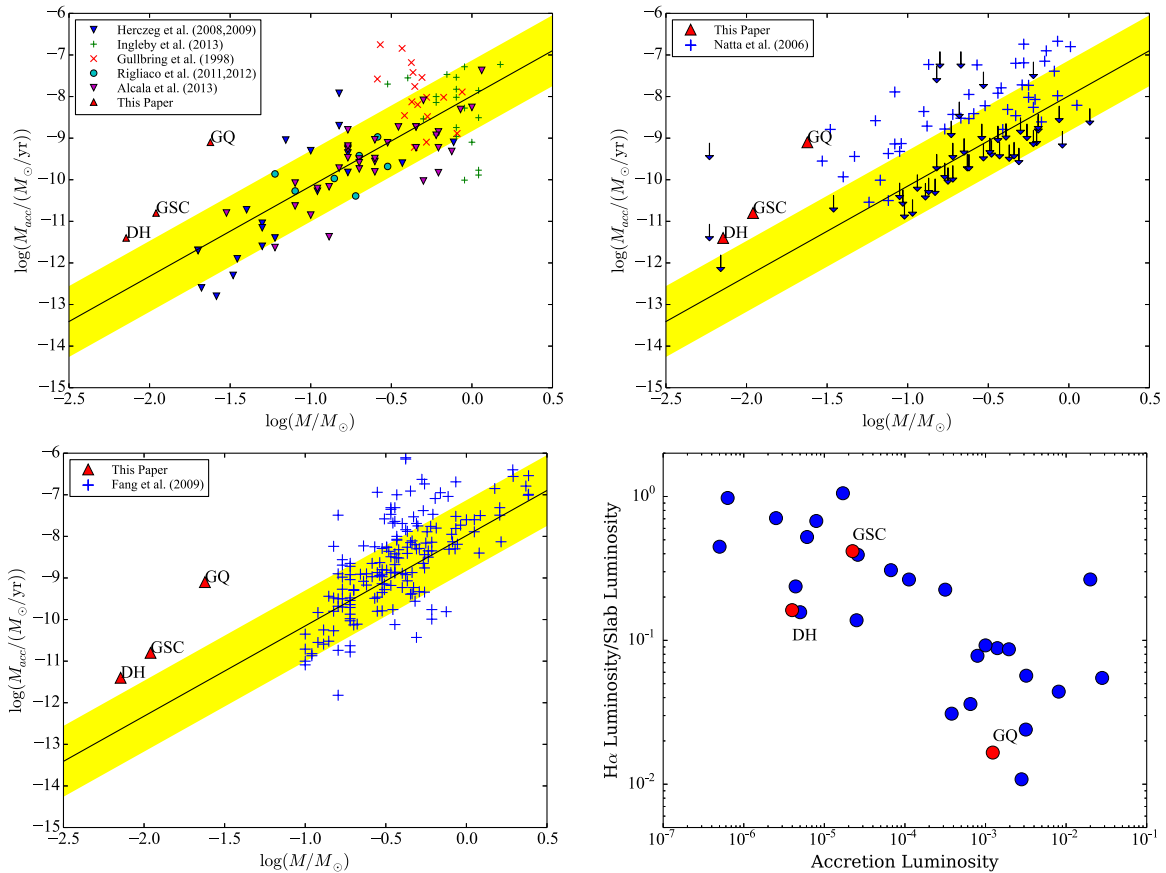


Figure 3. Accretion properties of the planetary mass companions compared to other young accretors. The three planetary mass companions have accretion rates that are one order of magnitude higher than expected from the correlation between object mass and UV excess accretion rates (upper left) and from accretion rates obtained from combined measurements of the $H\alpha$, $H\beta$, and $He\ I\ \lambda 5876$ lines (lower left), but are consistent with accretion rates measured in ρ Oph from $Pa\beta$ and $Br\gamma$ (upper right). The yellow shaded region shows the $\pm 1\sigma$ region from best fit line between mass and UV excess accretion rates and demonstrates the offsets between different studies. Lower right: the ratio of $H\alpha$ luminosity to slab luminosity increases with lower accretion luminosity.

for GQ Lup, possible J -band veiling from the accretion flow. Accretion luminosities and rates are similar for both approaches and are listed separately in Table 1.

The total accretion luminosity, L_{acc} , is calculated by adding excess continuum emission from the slab, L_{slab} , and the excess line emission, in this case $H\alpha$, $L_{H\alpha}$. The contribution of line emission to accretion rate measurements has been ignored in all previous publications. Following Gullbring et al. (1998), the accretion rate \dot{M} is obtained using

$$\dot{M} \approx \frac{1.25 R_* L_{acc}}{GM_*}. \quad (1)$$

The accretion luminosities and rates are listed in Table 1. The extinction uncertainty, here assessed as ~ 0.5 mag, dominates the factor of ~ 2 uncertainty in accretion luminosity. The factor of 3–5 uncertainty in accretion rate includes uncertainties in the radius and temperature. A detailed description of errors is discussed in Herczeg & Hillenbrand (2008). The uncertainty in extinction measurements introduces large errors in accretion luminosity and rate. For GQ Lup b, an assumed extinction of $A_V = 0.4$ mag would yield $\log(L_{acc}/L_\odot) = -3.7$ and $\log(\dot{M}/(M_\odot\text{ yr}^{-1})) = -10.1$.

4. DISCUSSION

Accretion is detected in both the $H\alpha$ line and in excess optical continuum emission from DH Tau b, GSC 06214–00210 b, and GQ Lup b. Accretion had been previously detected for all three

objects in $Pa\beta$ emission, although this emission was undetected in some spectra of GQ Lup b. For GSC 06214–00210 b, the accretion luminosity measured here (when including $H\alpha$ luminosity) of $\log L_{acc}/L_\odot = -4.6 \pm 0.5$ is similar to the -4.4 ± 1.3 measured from $Pa\beta$ by Bowler et al. (2011), although our direct measurements have much smaller error bars.

The accretion rates calculated here demonstrate that these objects have their own disks. The separations (100–300 AU) of the companion objects to their primaries are too large to support such high accretion rates. Moreover, there is no evidence to indicate the presence of a primordial disk around the primary star GSC 06214–00210. Disks around substellar secondaries have also been seen in transit in front of a primary star in the 1SWASP J1407 system (Mamajek et al. 2012).

These sources are at the brown dwarf–planet mass boundary, are among the lowest mass sources with accretion measured directly from excess optical spectra, and are the lowest mass companions to stellar primary stars with measured accretion rates. Previous detections of accretion onto very low mass objects has been diagnosed through emission lines (e.g., Bowler et al. 2011; Joergens et al. 2013), which are indirect probes of accretion rate, and directly from the Balmer jump for several free-floating brown dwarfs (e.g., Herczeg et al. 2009; Rigliaco et al. 2012).

Figure 3 shows correlations between mass and accretion rate, as measured directly from excess Balmer continuum emission and indirectly from line luminosities. The accretion rates measured here are more than one order of magnitude

higher than expected from the correlations obtained for accretion rates measured directly from excess Balmer continuum emission (Gullbring et al. 1998; Herczeg & Hillenbrand 2008; Herczeg et al. 2009; Rigliaco et al. 2011, 2012; Ingleby et al. 2013; Alcalá et al. 2014) and separately for accretion rates measured indirectly from optical lines (Fang et al. 2009). This offset is not detected when comparing our accretion rates to those measured for stars in ρ Oph from Pa β and Br γ (Natta et al. 2006). However, the ρ Oph accretion rates at low object mass are anomalous relative to other studies, either because the objects are younger or because of methodological differences. Moreover, half of the stars in ρ Oph with disks have only upper limits in accretion luminosity. While differences in ages can complicate these comparisons, the high accretion rates are consistent with literature values only if the objects have the same age as ρ Oph and are therefore younger than the stars in their parent associations. Although upper limits of our non-detections should be considered, our program was complete for every known planetary mass companion at the time of proposal submission and these three objects are all high outliers in accretion rate.

These results suggest that wide planetary mass companions have higher accretion rates than expected for objects formed from the fragmentation of a protostellar core. If correct, then models for their formation and evolution models for the formation and evolution should account for these high accretion rates and the survival of their disks. Since core accretion is unlikely to form planets at such large radii in the disk, the systematically high accretion rates suggests either formation by gravitational instability or different disk evolution between single planetary mass objects/brown dwarfs versus those around stellar mass objects.

In previous studies, the accretion rate has been calculated from only the luminosity from the continuum emission. The H α luminosity is usually negligible for solar mass stars. For GQ Lup b and DH Tau b, the H α luminosity is 5% of the accretion continuum luminosity. However, for GSC 06214–0210 b, H α luminosity is equivalent to the total slab luminosity and accounts for 50% of the total accretion luminosity. Figure 3 shows the ratio of H α luminosity to accretion continuum luminosity from our work and from literature values (Herczeg & Hillenbrand 2008; Herczeg et al. 2009; Rigliaco et al. 2011, 2012), when measured simultaneously. The percentage of accretion luminosity that escapes in H α increases with lower accretion luminosities, possibly because of lower opacities and temperatures in the accretion shock and accretion funnel flow. Especially for low-mass objects with high $L_{H\alpha}/L_{\text{slab}}$ ratios, the accretion luminosities may be underestimated if Ly α is significantly brighter than H α . Unfortunately, Ly α emission from CTTs is difficult to observe because of line-of-sight H I absorption (Schindhelm et al. 2012). Indeed, the one high point in $L_{H\alpha}/L_{\text{acc}}$ among solar mass stars is TW Hya, which has an Ly α luminosity of $\sim 0.01 L_{\odot}$, or $0.5 L_{\text{slab}}$ and $2 L_{H\alpha}$ (Herczeg et al. 2004).

The strength of H α emission suggests an alternate method to search for forming protoplanets. The accretion continuum is spread over a large wavelength region, so the contrast between the primary and an accreting secondary star is still high. However, if 10%–50% of the accreting flux is produced in a single line, then the contrast between the star and any accreting companion becomes much smaller. Targeted searches for H α emission from companions may be a powerful technique to

find proto-Jupiters that are undergoing their main phase of gas accretion. Indeed, Close et al. (2014) recently used the new Magellan Adaptive Optics system to detect H α emission and calculate accretion onto the faint companion around HD 142527.

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