



UNIVERSITY OF ILLINOIS
URBANA

AERONOMY REPORT NO. 55

ANALYSIS OF PARTIAL-REFLECTION DATA FROM THE SOLAR ECLIPSE OF JULY 10, 1972

(NASA-CR-136123) ANALYSIS OF
PARTIAL-REFLECTION DATA FROM THE SOLAR
ECLIPSE OF 10 JUL. 1972 (Illinois Univ.)

N74-12480

128 p HC \$8.50

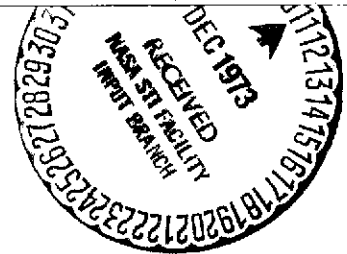
CSCL 03A

Unclas

G3/30 15935

129

by
T. A. Bean
S. A. Bowhill



October 1, 1973

Supported by
National Aeronautics and Space Administration
Grant NGR 14-005-181

Aeronomy Laboratory
Department of Electrical Engineering
University of Illinois
Urbana, Illinois

CITATION POLICY

The material contained in this report is preliminary information circulated rapidly in the interest of prompt interchange of scientific information and may be later revised on publication in accepted aeronomic journals. It would therefore be appreciated if persons wishing to cite work contained herein would first contact the authors to ascertain if the relevant material is part of a paper published or in process.

A E R O N O M Y R E P O R T

N O. 55

ANALYSIS OF PARTIAL-REFLECTION DATA FROM
THE SOLAR ECLIPSE OF JULY 10, 1972

by

T. A. Bean
S. A. Bowhill

October 1, 1973

Supported by
National Aeronautics and Space Administration
Grant NGR 14-005-181

Aeronomy Laboratory
Department of Electrical Engineering
University of Illinois
Urbana, Illinois

PRECEDING PAGE BLANK NOT FILMED

ABSTRACT

Partial-reflection data collected for the eclipse of July 10, 1972 as well as for July 9 and 11, 1972, are analyzed to determine eclipse effects on *D*-region electron densities. The partial-reflection experiment was set up to collect data using an on-line PDP-15 computer and DECTape storage. Except for a couple of changes, the experiment was the same setup as used by *Birley and Sechrist* [1971]. The electron-density profiles show good agreement with results from other eclipses. The partial-reflection programs were changed after the eclipse data collection to improve the operation of the partial-reflection system. These changes were mainly due to expanded computer hardware and have simplified the operations of the system considerably.

TABLE OF CONTENTS

	Page
ABSTRACT	iii
TABLE OF CONTENTS.	iv
LIST OF FIGURES.	v
1. INTRODUCTION	1
2. PRODUCTION AND LOSS OF THE <i>D</i> -REGION IONIZATION	5
2.1 <i>Ionization Sources</i>	5
2.2 <i>Formation of Ions in the D Region</i>	8
2.3 <i>Recombination</i>	13
2.4 <i>Expected Results</i>	14
2.5 <i>Statement of the Problem</i>	15
3. EXPERIMENTAL TECHNIQUE	19
3.1 <i>Development of Receiving and Storing Data</i>	20
3.2 <i>Partial-Reflection Data Collection for the Solar Eclipse</i>	26
3.3 <i>Real-Time Data Storage and Automatic Processing</i>	37
3.4 <i>Noise Rejection</i>	43
3.5 <i>Converting A_x/A_0 Ratios to Electron-Density Profiles</i>	48
3.6 <i>Equipment Testing</i>	51
3.7 <i>Future Development</i>	54
4. EXPERIMENTAL RESULTS	55
4.1 <i>Reduction of Data</i>	55
4.2 <i>Electron-Density Results</i>	58
4.3 <i>Theoretical Applications</i>	64
4.4 <i>Summary</i>	71
5. CONCLUSIONS	77
5.1 <i>Review of Results</i>	77
5.2 <i>Suggestions for Further Work</i>	78
REFERENCES	81
APPENDIX	86

LIST OF FIGURES

Figure	Page
1.1	2
The obscuration function of visible light at a height of 75 km for the eclipse of July 10, 1972, near Urbana, Illinois	
1.2	3
Obscuration functions for visible light (V), Lyman alpha (L_{α}) ultraviolet (UV), and X-ray (X) ionizing fluxes for the 1966 solar eclipse from <i>Sears</i> [1972].	
2.1	7
Ion-pair production rates from various D-region ionization sources as given by <i>Sechrist</i> [1972]	
2.2	9
Average variations in 2-8 Å X-ray flux during which partial-reflection data were collected on July 9, 10, and 11, 1972	
2.3	10
Flow diagram of the formation of positive ions including conversion rates [<i>Donahue</i> , 1972]. Three-body rate constants are in units of $10^{-28} \text{ cm}^6 \text{ sec}^{-1}$; two-body rate constants not given by Donahue are from <i>Good, et al.</i> [1970].	
2.4	12
Block diagram [<i>Thomas</i> , 1971] showing the negative ion chemistry during the day. The lifetimes of electrons and each ion are for a height of 65 km	
2.5	16
The variation of the solar zenith angle for July 10, 1972. The partial-reflection data collection period is shown as well as the time of maximum obscuration for the eclipse.	

Figure	Page
2.6	Variation of electron density during a solar eclipse at March equinox, mid-day, and sunspot minimum, at middle latitudes [Deeks, 1966] 17
2.7	Electron-density profiles for the eclipse of July 20, 1963 [Smith, et al., 1965]. Profiles 1, 2, 3, and 4 refer to obscurations of 92%, 86%, 40%, and 2%, respectively. The solar zenith angle was 55° at totality and 61° at 40% obscuration 18
3.1	Block diagram of the partial-reflection transmitter. 21
3.2	Partial-reflection antenna arrays for the Aeronomy Field Station. 22
3.3	Block diagram of the partial-reflection system 23
3.4	Typical frame of data as collected by Henry [1966] 24
3.5	Block diagram of the revised receiver used to operate with a PDP-15 computer 28
3.6	The RF amplifier module for the receiver 29
3.7	The IF and DC amplifier module 30
3.8	The encode pulses as set up by Birley and Sechrist [1971] used to collect data during the eclipse, and the revised encode pulses used by the present programs 31
3.9	The encode pulse circuitry used to produce the former and present encode pulses. 32
3.10	Graphs of the input versus output of the receiver used for eclipse data collection (old receiver) and the receiver presently being used. The input and output values

Figure	Page
have been normalized to the maximum of the A/D converter (511).	34
3.11 The wiring diagram used to calibrate the receiver. The voltmeter is used in setting the initial signal level prior to calibrating	35
3.12 A diagram of the control flow of the partial-reflection programs. The programs operate on the API level of the preceding program unless otherwise stated. The collection and processing programs operated in parallel with the collection programs operating on API levels 5 and 6 whereas everything else operates serially.	41
3.13 A flow chart of the processing program PROC.	46
3.14 The collision frequencies used in the program CALC to obtain electron-density profiles.	50
4.1 Comparison of electron-density profiles on July 10 and 11, 1972. The data were taken at 1432 CST with the attenuator set at 30 dB.	56
4.2 Comparison of the A_x/A_o ratio at 72 km for July 9, 10, and 11	59
4.3 Comparison of the A_x/A_o ratio at 75 km for July 9, 10, and 11	60
4.4 Comparison of the A_x/A_o ratio at 78 km for July 9, 10, and 11	61
4.5 Median electron densities between 75 and 82.5 km	62
4.6 Median A_x/A_o profiles between 1400 and 1500 CST for each day	63

Figure	Page
4.7	Average electron densities for the three days using method two 65
4.8	Median electron densities between 67.5 and 75 km 66
4.9	Median electron-density profiles between 1400 and 1500 CST 67
4.10	Electron production rate between 75 and 82.5 km during the eclipse and during the control days. The NO distribu- tion used is from <i>Meira</i> [1971] 69
4.11	Theoretical electron densities between 75 and 82.5 km for eclipse and control day; calculated using an α_{eff} of 1.77×10^{-6} 70
4.12	The ratio of electron densities for the average of the control day as compared to the unobscured sun. 72
4.13	The graph of the ratios of the theoretical [e] for the unobscured sun to the experimental [e] for the eclipse as compared to the unobscured sun. The α_{eff} used for 75 to 82.5 km is 1.77×10^{-6} and for 78 to 87.5 km is 8.46×10^{-7} . . 73
4.14	Scatter plot correlating the electron density for July 9, 1972 between 75 and 82.5 km to the solar zenith angle. 74
4.15	The graph of the ratio of theoretical electron densities to the experimental electron densities for July 9 and 10, 1972 75
5.1	Rocket measurements of the positive-ion chemistry by <i>Krankowsky, et al.</i> [1972]. 79

1. INTRODUCTION

A solar eclipse can be thought of as the obscuration of solar radiation by the intervention of the moon between the sun and a point on the earth. This obscuring of the sun is a function of time which varies with the location on the earth, altitude, and the type of radiation. Depending on the wavelength of solar radiation and the ionospheric constituents, solar radiation can cause three chemical processes known as dissociation, ionization and excitation [Whitten and Poppoff, 1971]. The variation in solar radiation with time during a solar eclipse is given as an obscuration function and varies according to the different wavelengths of radiation. The obscuration function for visible light is easily calculated, being just that for the visible disk. Figure 1.1 shows this obscuration function for the eclipse of July 10, 1972 at 75 km altitude above the University of Illinois Aeronomy Field Station located near Urbana. At this location the eclipse was partial, with about 60% maximum obscuration. The obscuration functions for various other radiations during a total eclipse are shown in Figure 1.2 [Sears, 1972]. Notice the large difference between the obscuration functions for ultraviolet radiation and X-rays.

Solar radiation with wavelengths less than 2900Å causes various chemical reactions in the ionosphere [Whitten and Poppoff, 1971] with the most pronounced effects occurring in the D-region (50 to 90 km). For example, Turco and Sechrist [1970] show two orders of magnitude change in the electron density and more than three orders of magnitude change in CO_3^- and CO_4^- at 75 km during sunrise. Certain solar radiations greatly enhance the concentration of positive ions as well as the electron density so that during the daytime, except for during enhanced particle precipitation [Lauter and Knuth, 1967], the main ionization source above 70 km

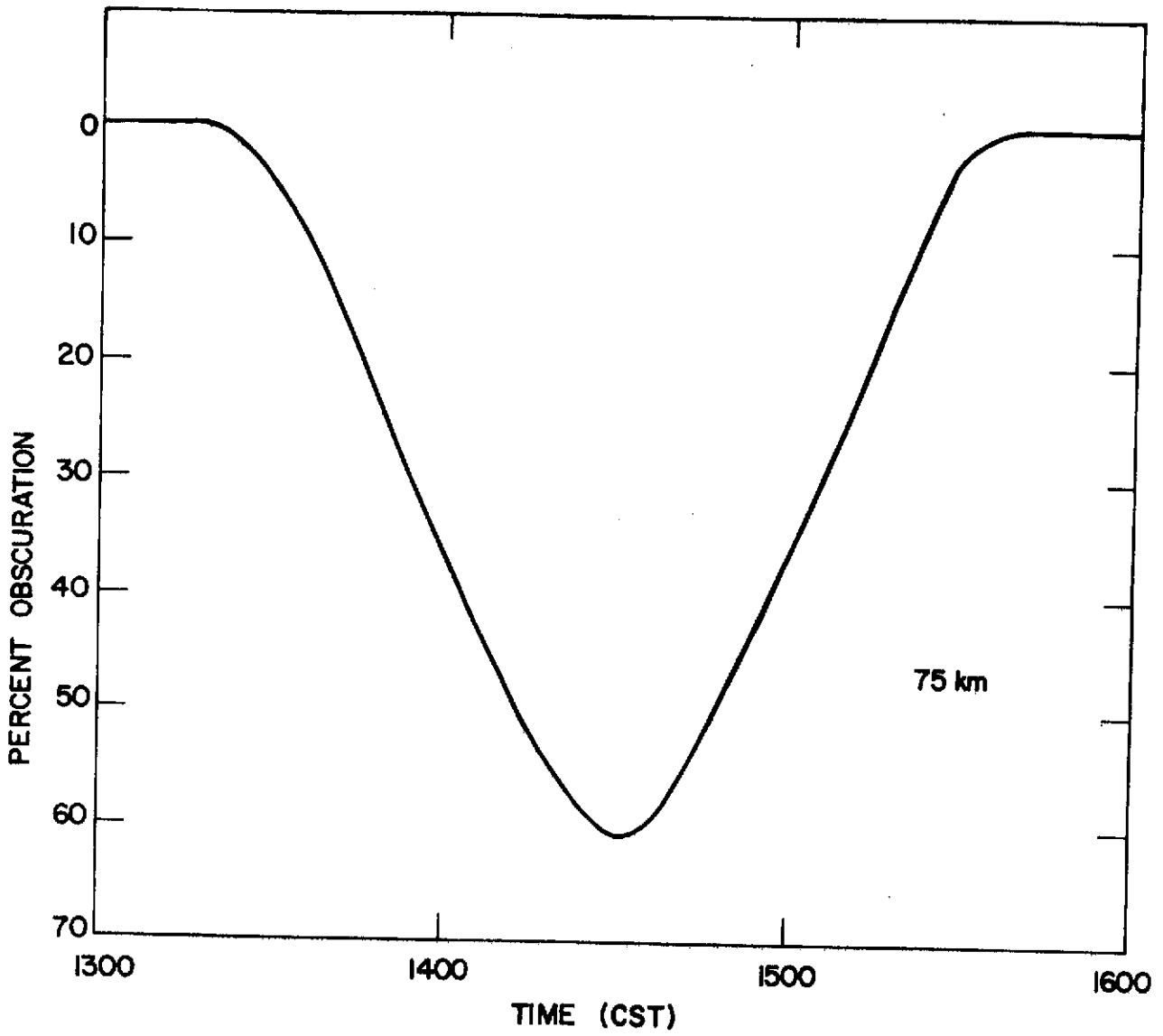


Figure 1.1 The obscuration function of visible light at a height of 75 km for the eclipse of July 10, 1972, near Urbana, Illinois.

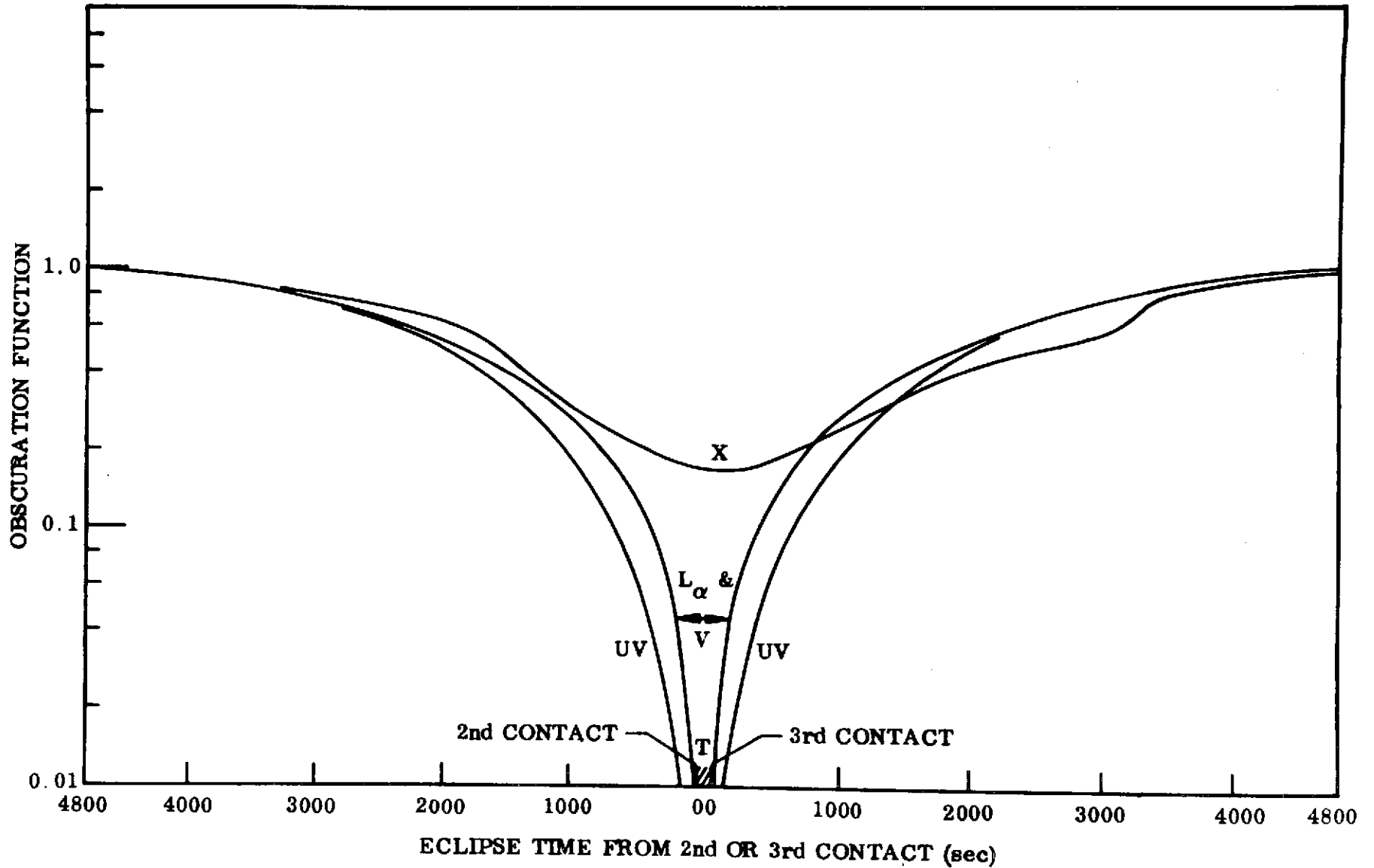


Figure 1.2 Obscuration functions for visible light (V), Lyman alpha (L_{α}), ultraviolet (UV), and X-ray (X) ionizing fluxes for the 1966 solar eclipse from *Sears* [1972].

is solar radiation as given in Section 2.1. Therefore by correlating the electron densities with the obscuration function for the ionizing radiation in a solar eclipse, values for the production and loss of positive ions and confirmation of the ionizing sources can be obtained.

Data from the *D* region have been obtained by both rocket measurements and ground-based techniques. Although rocket measurements seem to be more accurate [Mechtly, et al., 1967], the amount of data is limited by cost. Ground-based techniques can be set up anywhere and can gather large amounts of data, although the accuracy is not as great, and they are primarily limited to evaluating electron densities. One type of ground-based technique which is discussed in this paper is called the partial-reflection experiment. Data are collected using vertical incident radio waves which are partially reflected from the *D* region. The information obtained can be in one of two forms: differential absorption [Pirnat and Bowhill, 1968] and differential phase [Wiersma and Sechrist, 1972]. Partial-reflection data using the differential absorption method were collected from 1200 to 1700 CST for the solar eclipse of July 10, 1972, as well as July 9 and 11 as control days. The experiment was set up as described by Birley and Sechrist [1971] with two exceptions as described in Chapter 3. The solar and ionospheric conditions for this experiment are given in Chapter 2.

2. PRODUCTION AND LOSS OF THE D-REGION IONIZATION

Recently several papers have summarized the knowledge of the *D* region of the ionosphere. *Thomas* [1971] presents an overall review of the *D* region while theoretical models of the *D* region are presented by *Sechrist* [1972], *Ferguson* [1971], *Donahue* [1972], and *Radicezza and Stowe* [1970]. The *D* region is perhaps the most complicated part of the ionosphere as well as the most difficult part from which to obtain accurate data. The chemical composition is dependent on height and solar zenith angle [*Thomas*, 1971]; although it consists of neutral constituents, positive ions, negative ions, and free electrons, this chapter is mainly concerned with the processes of formation and loss of free electrons during the daytime (solar zenith angles less than 90°) and during a solar eclipse. Using results obtained from measurements on other eclipses, the expected results from the partial-reflection experiment are given.

2.1 Ionization Sources

Although there is general agreement on what ionizes the neutral *D*-region constituents, there is some doubt as to the relative importance of each source. The ionization sources for the daytime *D* region at midlatitudes, as given by *Mitra and Rowe* [1972] and by *Aikin* [1972] are:

- 1) Lyman- α (1216Å) ionizing nitric oxide (NO)
- 2) 1-8Å X-rays ionizing all constituents
- 3) 1027-1118Å ultraviolet radiation ionizing metastable $O_2(^1\Delta_g)$
- 4) Galactic cosmic rays ionizing all constituents.

Along with these sources precipitating electrons may be considered another source of free electrons, but is of prime importance only in the polar regions, at night, or after a magnetic storm [*Lauter and Knuth*, 1967] and will not be considered in this paper.

The primary ionization source below 70 km is considered to be galactic cosmic rays [Sechrist, 1972], although its effect may extend as high as 75 km [Keneshea, 1967]. The primary ionization source above 70 km is either (1) or (3) depending upon the nitric oxide distribution adopted. Few measurements of nitric oxide have been made, so most distributions available are from theoretical models. Distributions measured by Barth [1966] and Pearce [1969] are at least an order of magnitude greater than distributions calculated from theoretical models of the ionosphere [Mitra, 1966], but distributions measured by Meira [1971] below 85 km are about the same as those calculated by Shimazaki and Laird [1970]. Using distribution by Barth [1966] for NO, the primary ionization source between 70 and 80 km is Lyman- α ionizing NO, but using nitric oxide distributions given by Shimazaki and Laird [1972] and photoionization rates for $O_2(^1\Delta_g)$ given by Hunten and McElroy [1968], the main ionization source between 70 and 80 km is 1027-1118 Å UV radiation ionizing $O_2(^1\Delta_g)$ [Thomas, 1971]. Somoyajulu and Avadbanulu [1972] pointed out that according to measurements by Huffman, et al. [1971], photoionization of $O_2(^1\Delta_g)$ is important only above 80 km making Lyman- α the main ionization source. Figure 2.1 from Sechrist [1972] shows ion-pair production rates for various radiation during solar minimum. In any case the distribution of NO is important to the rate of production of free electrons between 70 and 80 km, and the distribution by Meira [1971] is used in this paper.

The variation of ionization sources (1) and (3) with respect to solar activity is small [Thomas, [1971], but 2-8 Å X-ray flux can change by several orders of magnitude. Typical X-ray fluxes for different solar activity as given by Aikin [1972] are less than 4×10^{-3} ergs cm^{-2} sec^{-1} for a quiet sun, between 4×10^{-4} and 4×10^{-3} ergs cm^{-2} sec^{-1} for moderate sun, and greater than 4×10^{-3} ergs cm^{-2} sec^{-1} for an active sun. A solar flare on July 11 at

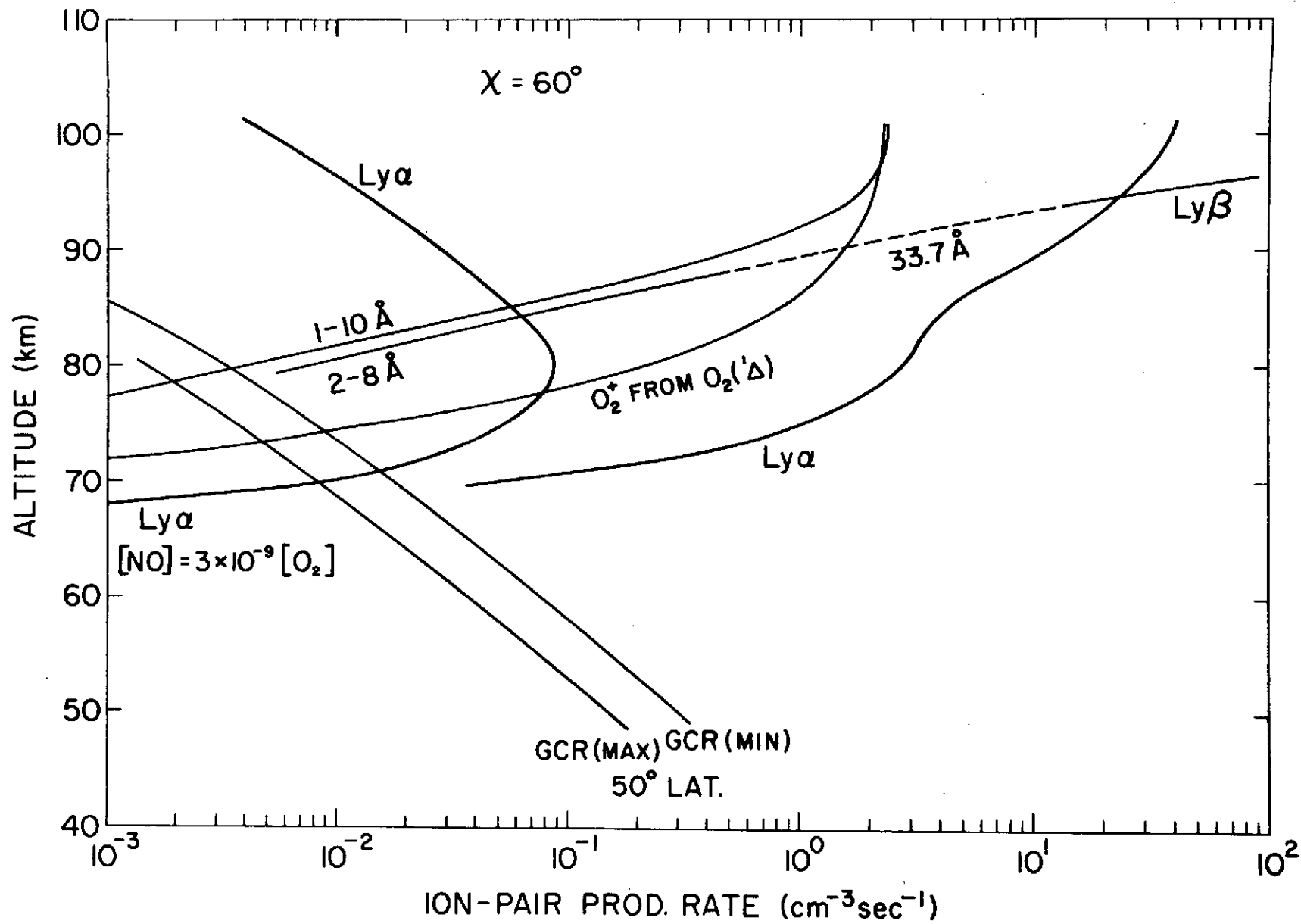
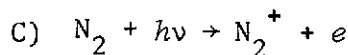
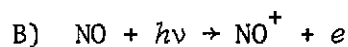
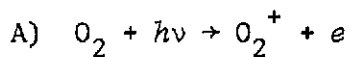


Figure 2.1 Ion-pair production rates from various D-region ionization sources as given by Sechrist [1972].

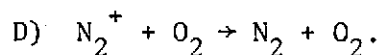
8:10 AM CST produced a 2-8 Å X-ray flux of 1.5×10^{-2} ergs cm^{-2} sec^{-1} . With an active sun or a solar flare 2-8 Å X-ray ionization can become the primary source of ionization. The 2-8 Å flux for July 9, 10, 11 in Figure 2.2 from the Solar Geophysical Data, 1973 (U. S. Department of Commerce) shows the solar activity to be quiet to moderate. The X-ray flux is expected to have little or no correlation with the electron density of the upper D region except for the X-ray burst near 1435 on July 11.

2.2 Formation of Ions in the D Region

The electron density between 70 and 85 km is dependent on the formation of positive ions. The three main ionization reactions for this region are:



as seen in Figure 2.3 adapted from *Mitra and Rowe* [1972] and *Donahue* [1972], which is a block diagram of the positive-ion chemistry at 75 km. The main loss process for N_2^+ is by the charge-exchange reaction:



This reaction is very fast (1×10^{-10} cm^3 sec^{-1}) [*Fehsenfeld, et al.*, 1965]. Therefore concentrations of N_2^+ are small and the production of O_2^+ is either by photoionization or by charge transfer. Electron production, therefore can be determined by the production of NO^+ and O_2^+ minus the formation of NO^+ by charge exchange reactions shown in Figure 2.3. Since the production of NO^+ is dependent on NO distributions, the production rate of the free electrons also depends on the NO distribution which can differ by at least an order of magnitude (Section 2.1).

The main positive ions between 70 and 80 km are hydrated ions of the form $\text{H}^+(\text{H}_2\text{O})_n$, n being some number greater than zero [*Narcisi and Bailey*, 1965].

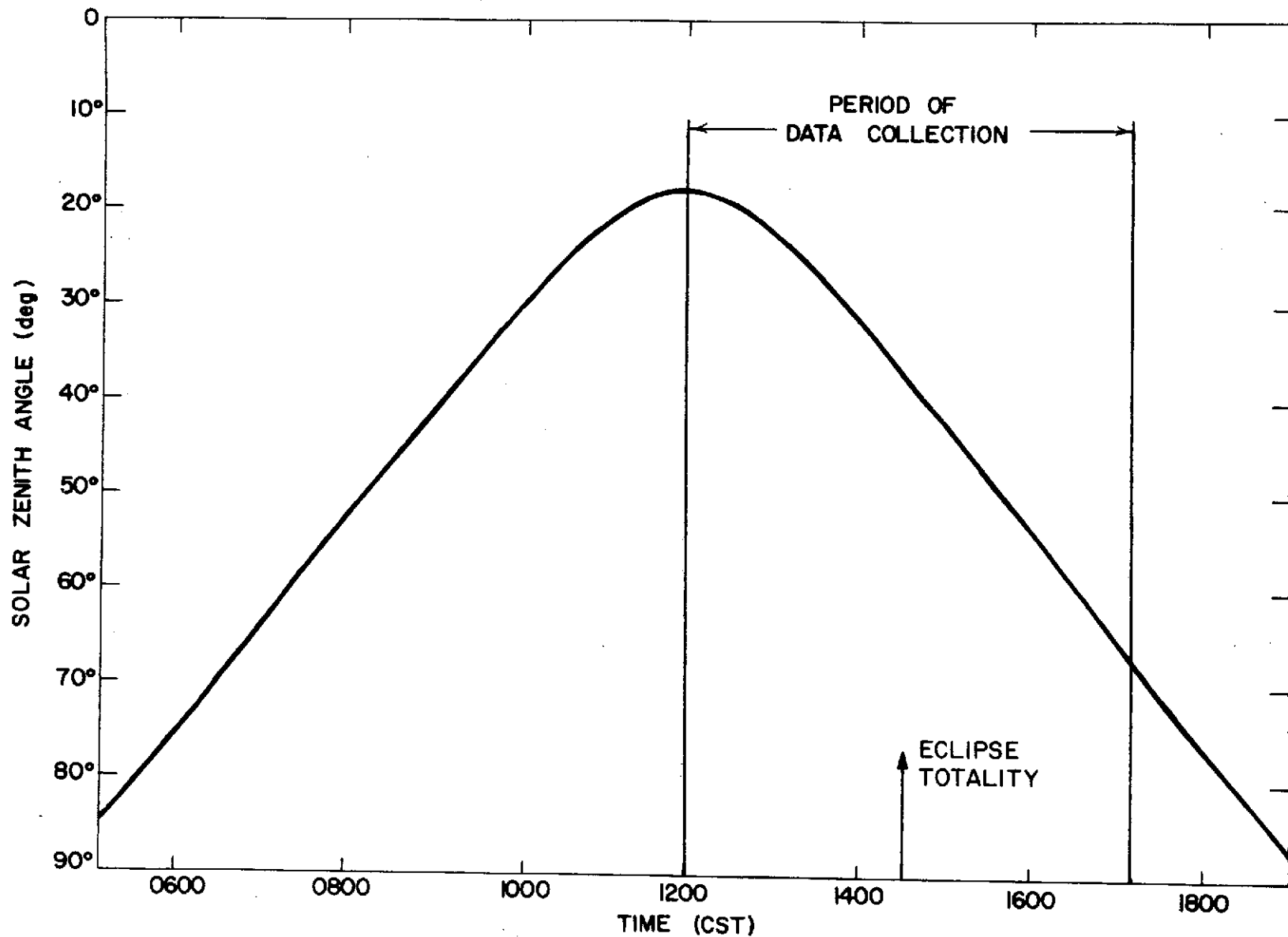


Figure 2.2 Average variations in 2-8 Å X-ray flux during which partial-reflection data were collected on July 9, 10, and 11, 1972.

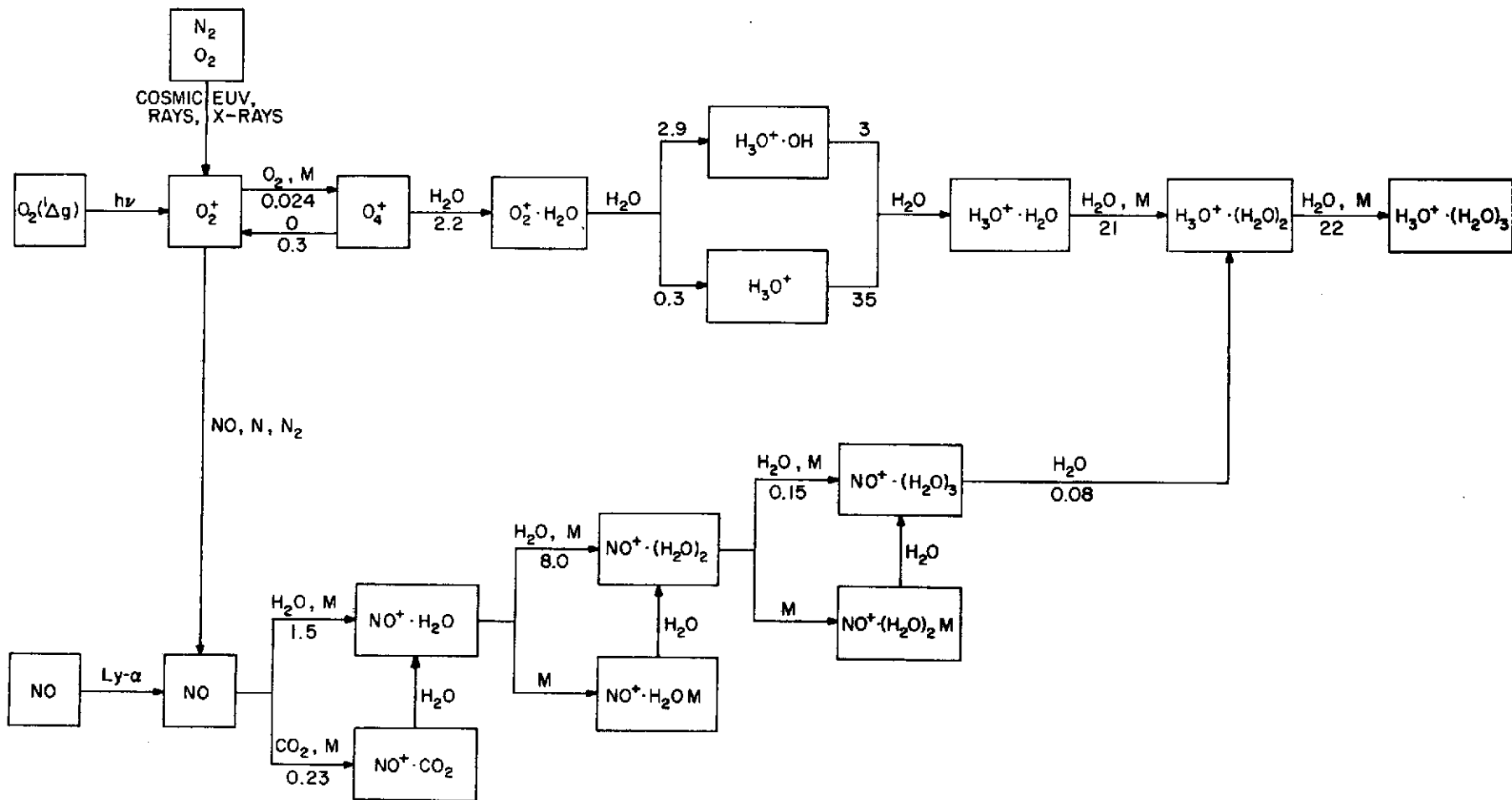


Figure 2.3 Flow diagram of the formation of positive ions including conversion rates [Donahue, 1972]. Three-body rate constants are in units of $10^{-28} \text{ cm}^6 \text{ sec}^{-1}$; two-body rate constants are in units of $10^{-9} \text{ cm}^3 \text{ sec}^{-1}$. Rate constants not given by Donahue are from Good, et al. [1970].

Two basic reaction schemes for the formation of water cluster ions as presented by *Fehsenfeld and Ferguson* [1969] are from NO^+ and beginning with the reaction $\text{O}_2^+ + \text{O}_2 + M \rightarrow \text{O}_4^+ + M$ where M is a third body. Both schemes are given in Figure 2.3. Each scheme raised several questions which are dealt with by *Donahue* [1972]. According to Figure 2.3, NO^+ creates hydrates with masses of 55 and higher, yet 19^+ and 37^+ are the dominant hydrates detected. Also the first three reactions with NO^+ are too slow relative to the loss rate. Problems with the O_2^+ scheme are: it seems to ignore the large NO^+ concentration and the ionization of $\text{O}_2(^1\Delta_g)$ seems to be an overestimation according to *Huffman, et al.* [1971], but this may be the main source of water clusters between 77 and 85 km [*Donahue, 1972*]. Even with the large number of hydrated ions, the rapid recombination rate competes with the formation of hydrated ions [*Thomas, 1971*]. This recombination represents the main loss process for free electrons between 70 and 80 km.

The formation of negative ions would constitute a loss of free electrons by the attachment reaction;

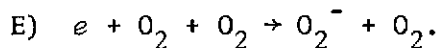
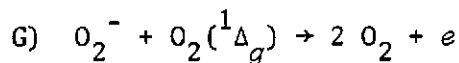
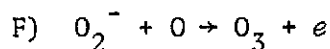


Figure 2.4 by *Thomas* [1971], giving a scheme for the daytime negative electrons at 65 km, shows reaction (E) to be fast, but the loss reactions



are much faster. Although the formation of O_4^- is rapid, there is rapid return to O_2^- . The negative ion chemistry is dependent on atomic oxygen and $\text{O}_2(^1\Delta_g)$ concentrations. At night these concentrations decrease so that reaction (E) constitutes an important loss process for free electrons.

At eclipse totality free electron production is reduced to that comparable of nighttime electron production, and the production of atomic oxygen and metastable $\text{O}_2(^1\Delta_g)$ are also greatly reduced [*Shimazaki and Laird, 1972*]. By

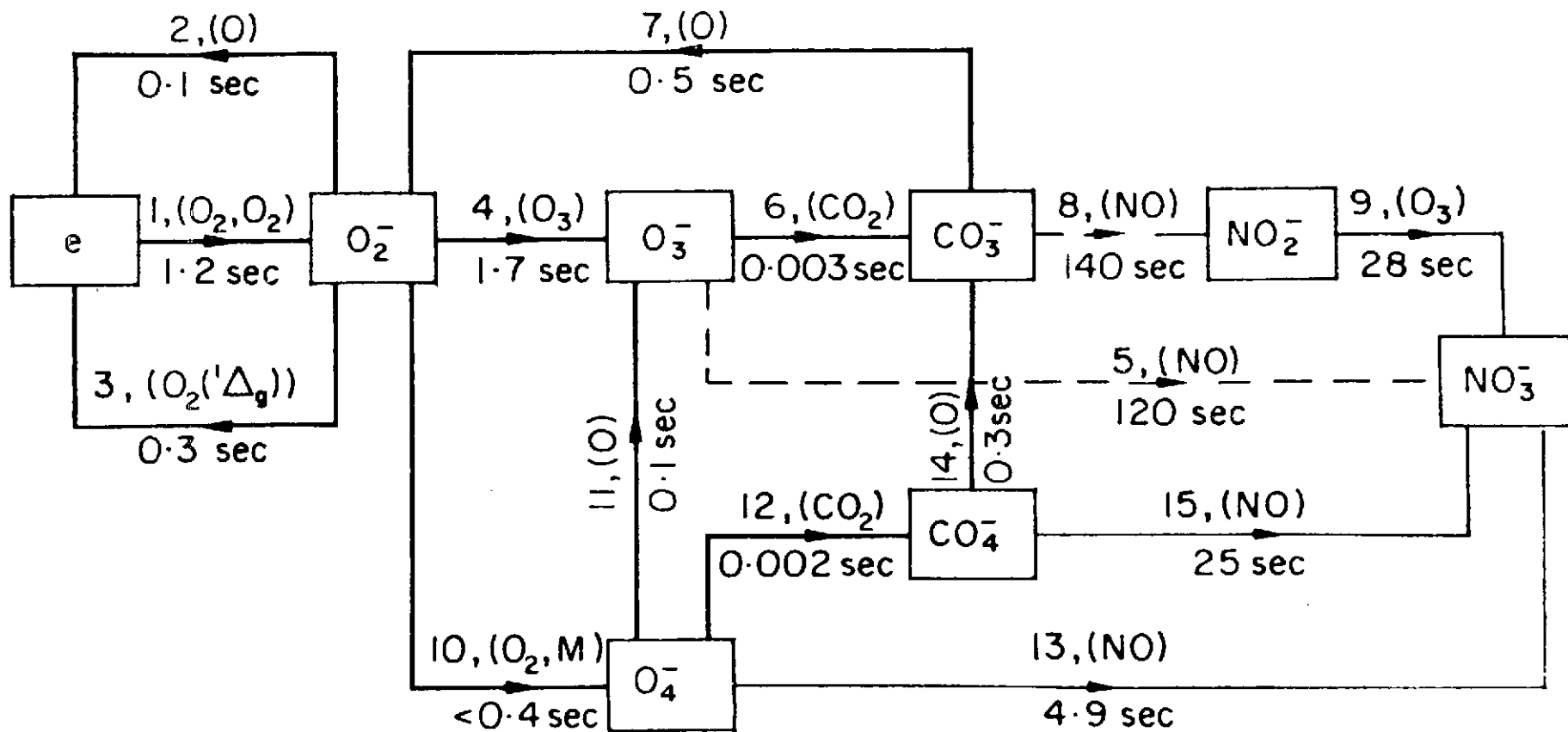


Figure 2.4 Block diagram [Thomas, 1971] showing the negative ion chemistry during the day. The lifetimes of electrons and each ion are for a height of 65 km.

comparison of eclipse data, *Mechtly, et al.* [1972] shows the possibility of attachment reactions as being the main loss process at totality. This would mean a large reduction in O and $O_2(^1\Delta_g)$, but the reduction measured by *Hunt* [1965] during an eclipse shows less than an order of magnitude change in atomic oxygen. More measurements of atomic oxygen are needed during eclipses to determine more accurately the loss process for free electrons during totality of a solar eclipse.

2.3 Recombination

Above 70 km during the daytime, negative-ion chemistry is not important; so the main loss process of free electrons above 70 km is by recombination with positive ions. The continuity equation for electrons as given by *Whitten and Poppoff* [1971] is:

$$\frac{d[e]}{dt} = \left(\frac{q}{1+\lambda}\right) - (\alpha_D + \lambda\alpha_i)[e]^2 - \left(\frac{[e]}{1+\lambda}\right) \frac{d\lambda}{dt} \quad (2.1)$$

where $[e]$ is the electron density, λ is the ratio of negative ion concentrations to electron densities, q is the ionization rate, α_D is the ion-electron recombination coefficient, and α_i is the ion-ion recombination coefficient.

With the assumption that variation in λ is insignificant, then $d\lambda/dt = 0$ and defining an effective recombination coefficient as $\alpha_{\text{eff}} = \alpha_D + \lambda\alpha_i$, Equation (2.1) reduced to:

$$\frac{d[e]}{dt} = \left(\frac{q}{1+\lambda}\right) - \alpha_{\text{eff}}[e]^2 \quad (2.2)$$

During a solar eclipse at totality, the electron production decreases by several orders of magnitude. Using an ionization rate of zero ($q = 0$), α_{eff} can be obtained from Equation (2.3) for short intervals of time.

$$\alpha_{\text{eff}} = \frac{\Delta[e]}{\Delta t} [e]^2 \quad (2.3)$$

With small changes in the electron density α_{eff} can be obtained by the approximation [Mitra and Rowe, 1972]

$$\alpha_{\text{eff}} = q/[e]^2(1 + \lambda) \quad (2.4)$$

Below 70 km the problem is complicated by the presence of negative ions [Mitra and Rowe, 1972] for which a time dependent analysis of the negative reaction scheme has to be used [Thomas, 1971]. As discussed in Section 2.2, there is the possibility of loss by attachment. Many problems about the loss process still remain unsolved including the question of the NO distribution.

2.4 Expected Results

Figure 1.2 by Sears [1972] gives the obscuration function for different D-region solar ionization sources from the eclipse of 1966. Lyman- α and visible light have the same obscuration function but not so with UV and X-rays. The obscuration function for visible light at Urbana, Illinois for July 10, 1972 (Figure 1.1) is therefore expected to be different from the obscuration function for ultraviolet radiation and X-rays. Using the maps of the sun given in Solar-Geophysical Data, 1972 (U.S. Department of Commerce) and the moon's movement across the sun's disk, an idea of the obscuration function for different solar radiations can be obtained. Since the solar activity during the eclipse was quiet to moderate, the predominate ionization source between 70 and 80 km is expected to be Lyman- α .

The total obscuration is about 60%, therefore data is used from previous eclipses with a similar obscuration and about the same solar zenith angle. The

solar zenith angle is shown in Figure 2.5 to be about 37° . Figure 2.6 by Deeks [1966] gives various electron densities for an eclipse during March equinox noon at sunspot minimum. Figure 2.7 by Smith, *et al.* [1965] gives electron density distributions for various obscurations of the eclipse of July 20, 1963. In Figure 2.6 the electron density for 60% obscuration shows little change until above 70 km. For Figure 2.7 at 40% obscuration the electron density at 75 km has no change while above and below this altitude show marked changes. Below 75 km the change is, therefore, expected to be no larger than above 75 km and the change is expected to be approximately 36% (from equation (2.4)). Due to the changing solar zenith angle, the magnitude of the slope of the changing electron densities before the maximum obscuration of the sun is expected to be greater than the slope after maximum obscuration.

2.5 *Statement of the Problem*

The purpose of this paper is to present the setting up, collection, and analysis of the partial-reflection data taken before, during and after a solar eclipse and to present changes made in the partial-reflection computer programs in order to simplify the operation and more effectively reject noise.

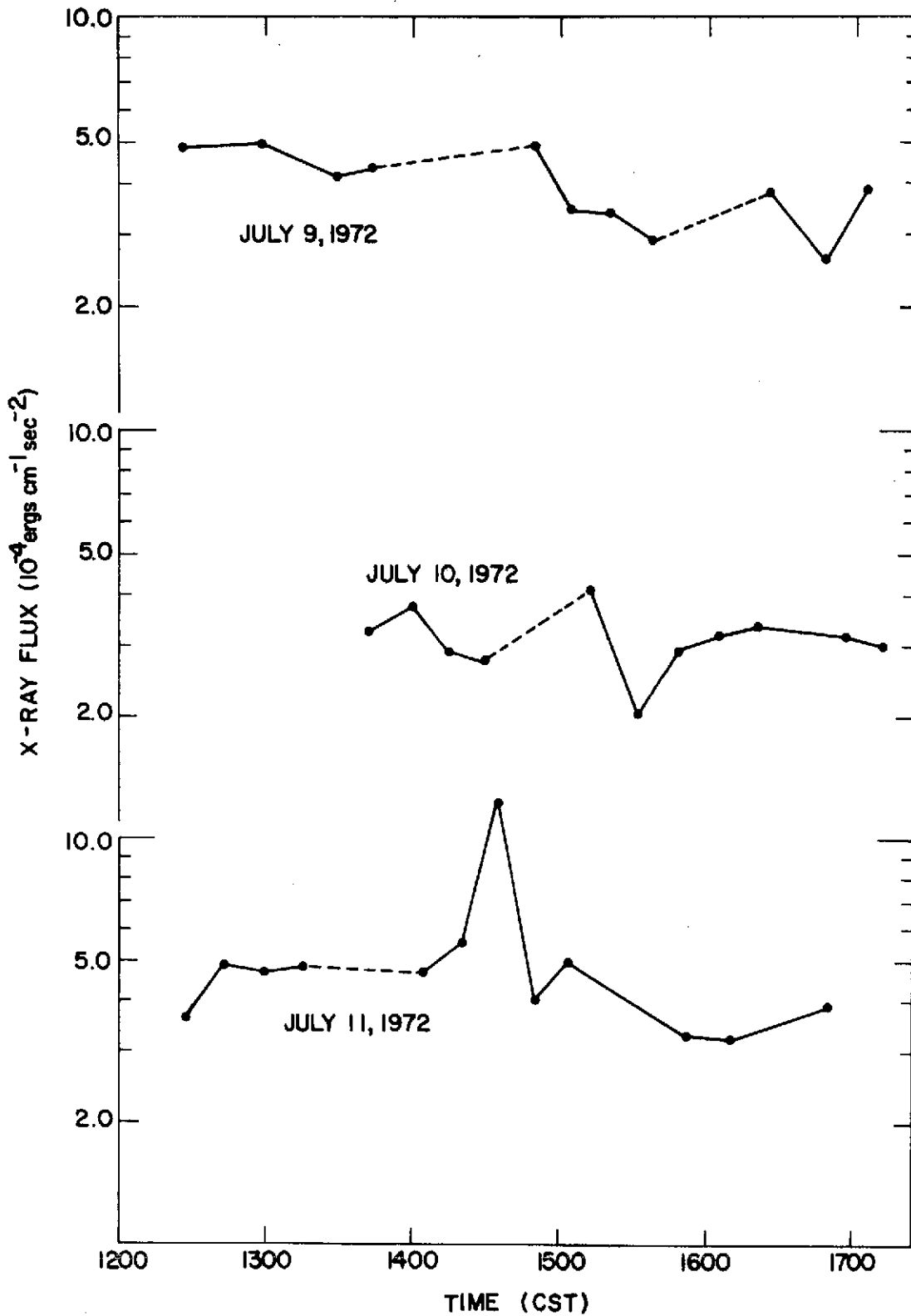


Figure 2.5 The variation of the solar zenith angle for July 10, 1972. The partial-reflection data collected period is shown as well as the time of maximum obscuration for the eclipse.

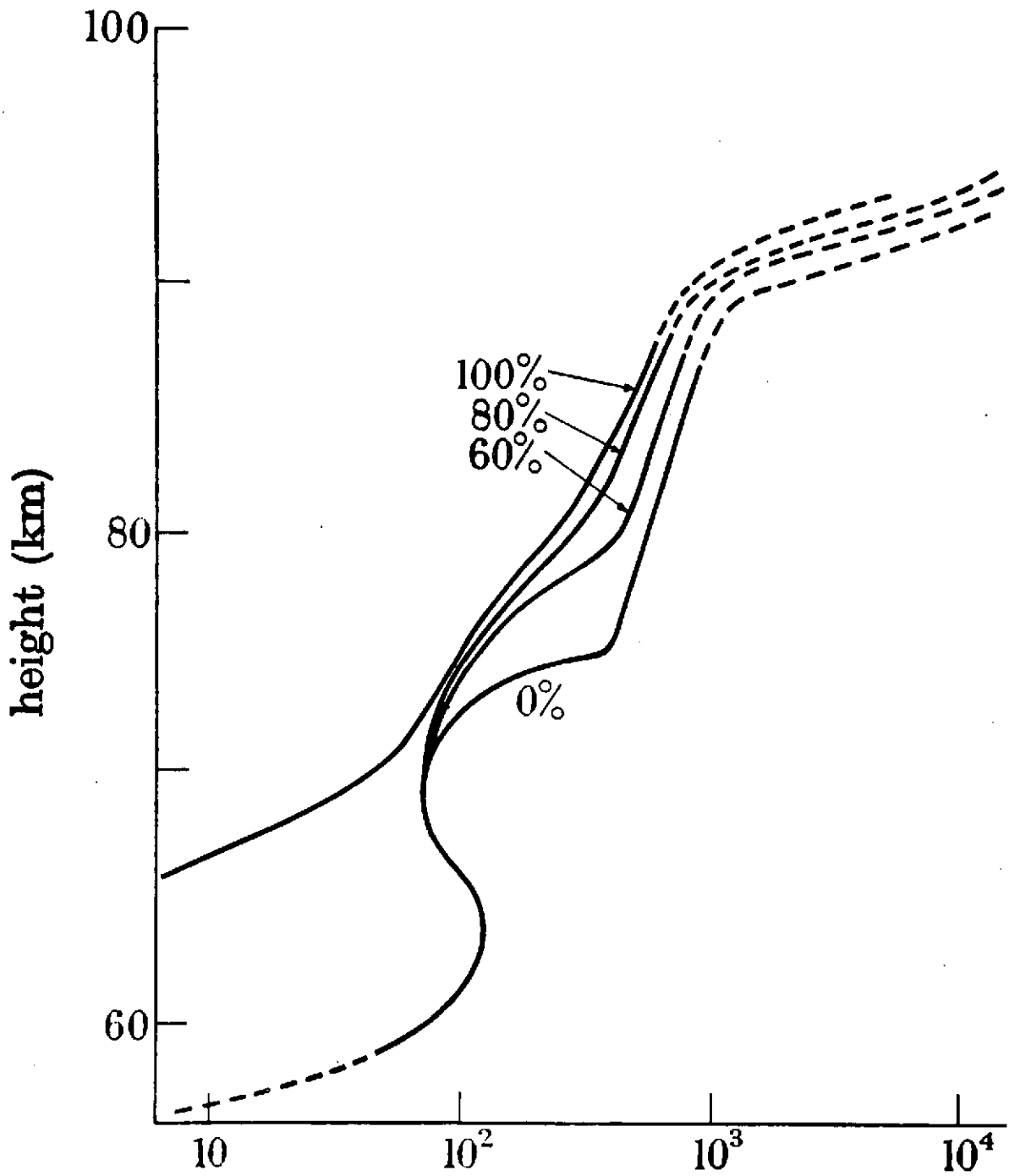


Figure 2.6 Variation of electron density during a solar eclipse at March equinox, mid-day, and sunspot minimum at middle latitudes [Deeks, 1966].

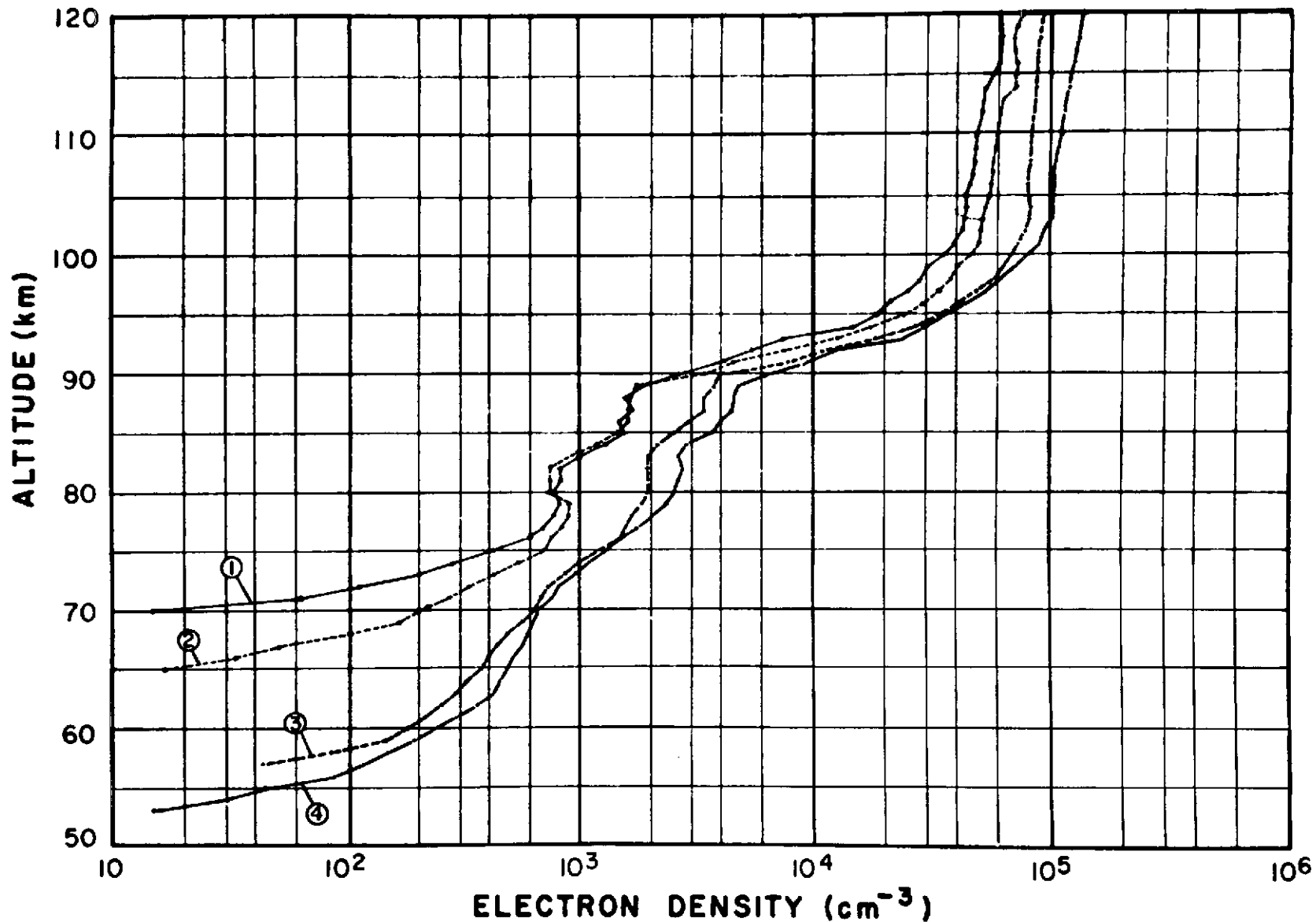


Figure 2.7 Electron-density profiles for the eclipse of July 20, 1963 [Smith, *et al.*, 1965]. Profiles 1, 2, 3, and 4 refer to obscurations of 92%, 86%, 40%, and 2%, respectively. The solar zenith angle was 55° at totality and 61° at 40% obscuration.

3. EXPERIMENTAL TECHNIQUE

The partial-reflection experiment was first performed by *Gardner and Pawsey* [1953]. Electron densities were deduced for 65 to 82 km from partially reflected, circularly polarized radio waves. The transmitter operated at 1 kw during each 30 μ sec pulse with a center frequency of 2.28 MHz, and the partially reflected signals were displayed on an A-scan oscilloscope. Several improvements have been made in the experiment and are discussed by *Pirnat and Bowhill* [1968].

Gregory [1956] used an increase in transmitter power of 4 kw and a decrease in the transmitter pulse width to 9 μ sec. These changes improved the amplitude and resolution of the partial reflections. *Fejer and Vice* [1959] developed an improved receiving and storing method using a dual-beam cathode-ray tube oscilloscope and camera. The system was operated at 1.83 and 2.63 MHz. *Belrose and Burke* [1964] also operated at two different frequencies (2.66 and 6.275 MHz) and transmitter power of 1 Mw, were able to obtain electron densities from the *D* and *E* region. *Belrose and Burke* [1964] were the first to use the generalized Appleton-Hartree formulas by *Sen and Wyller* [1960] for partial-reflection application.

Using the generalized Appleton-Hartree formulas and several approximations, the ratios of partially reflected extraordinary waves (A_x) to the partially reflected ordinary wave (A_o) for two heights can be used to calculate electron densities [*Pirnat and Bowhill*, 1968 and *Reynolds and Sechrist*, 1970]. The ratio A_x/A_o at each height is inversely related to the absorption by the expression $\exp(2\int_0^h k_x - k_o)$ from which the name differential absorption originates. At the University of Illinois the electron density was calculated directly from these ratios, and as seen in Chapter 4, small changes in these ratios can produce large variations in the electron densities.

Henry [1966] designed and built the hardware for the partial-reflection experiment at the University of Illinois. The transmitter that is presently being used was built for the purpose of making shipboard measurements. This transmitter operates at 40 kw during each 20 μ sec pulse and with 5 pulses per second. The center frequency is 2.66 MHz with a 50-ohm unbalanced output. Figure 3.1 shows a block diagram of the transmitter. The reduction of power from the initial 50 kw used is to give longer life to the tubes used, and the pulse is shortened from 50 μ sec used by *Henry* [1966] for better height resolution.

Figure 3.2 shows the two antenna arrays used to transmit and receive circularly polarized signals. Each array consists of 30 half-wave dipoles in the north-south direction and 30 in the east-west direction [*Wiersma and Sechrist*, 1972]. Each direction has matching networks that differ by 90° from the other direction of the same array to give a circularly polarized radio wave as shown in Figure 3.3. Each array gives approximately 22 dB gain with the main beam in the vertical direction. The first sidelobe is down 14 dB. Since both arrays are the same, this is a decrease of approximately 30 dB in the sidelobes relative to the main signal which has 44 dB gain. Further details on the antennas are given by *Pirmat and Bowhill* [1968] and *Reynolds and Sechrist* [1970].

3.1 *Development of Receiving and Storing Data*

The receiver, storage and timing controls have had two main changes in the development of the partial-reflection system. The experiment was originally set up using photographic film to store the partially reflected signals as displayed on an oscilloscope (see Figure 3.4). The controlling circuitry or pulser sent pulses of 30 volts to the transmitter, receiver, and camera. The pulser has remained the same with the exception of the addition of extra control circuitry depending on the storage method. The amplitudes of the received signals

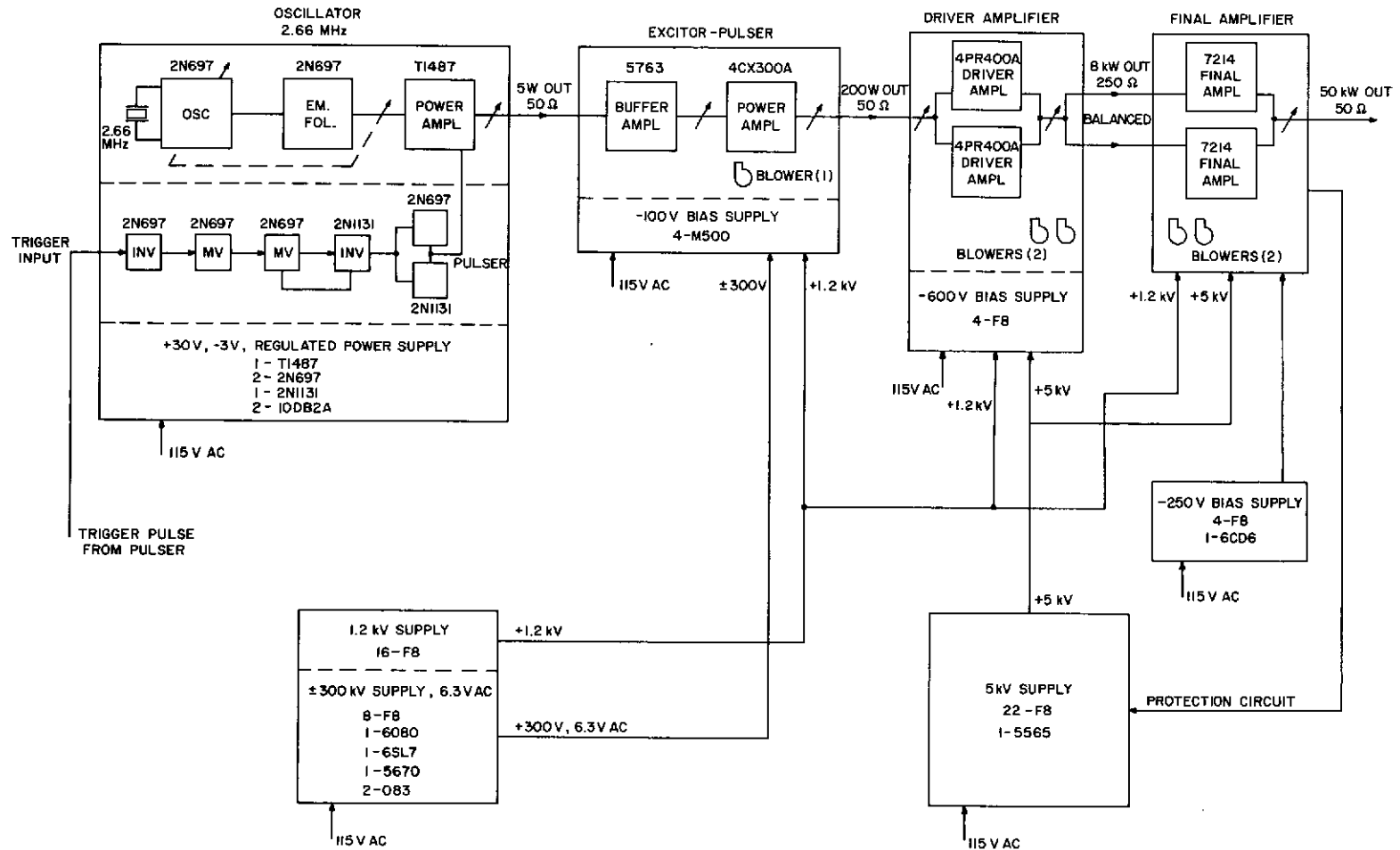


Figure 3.1 Block diagram of the partial-reflection transmitter.

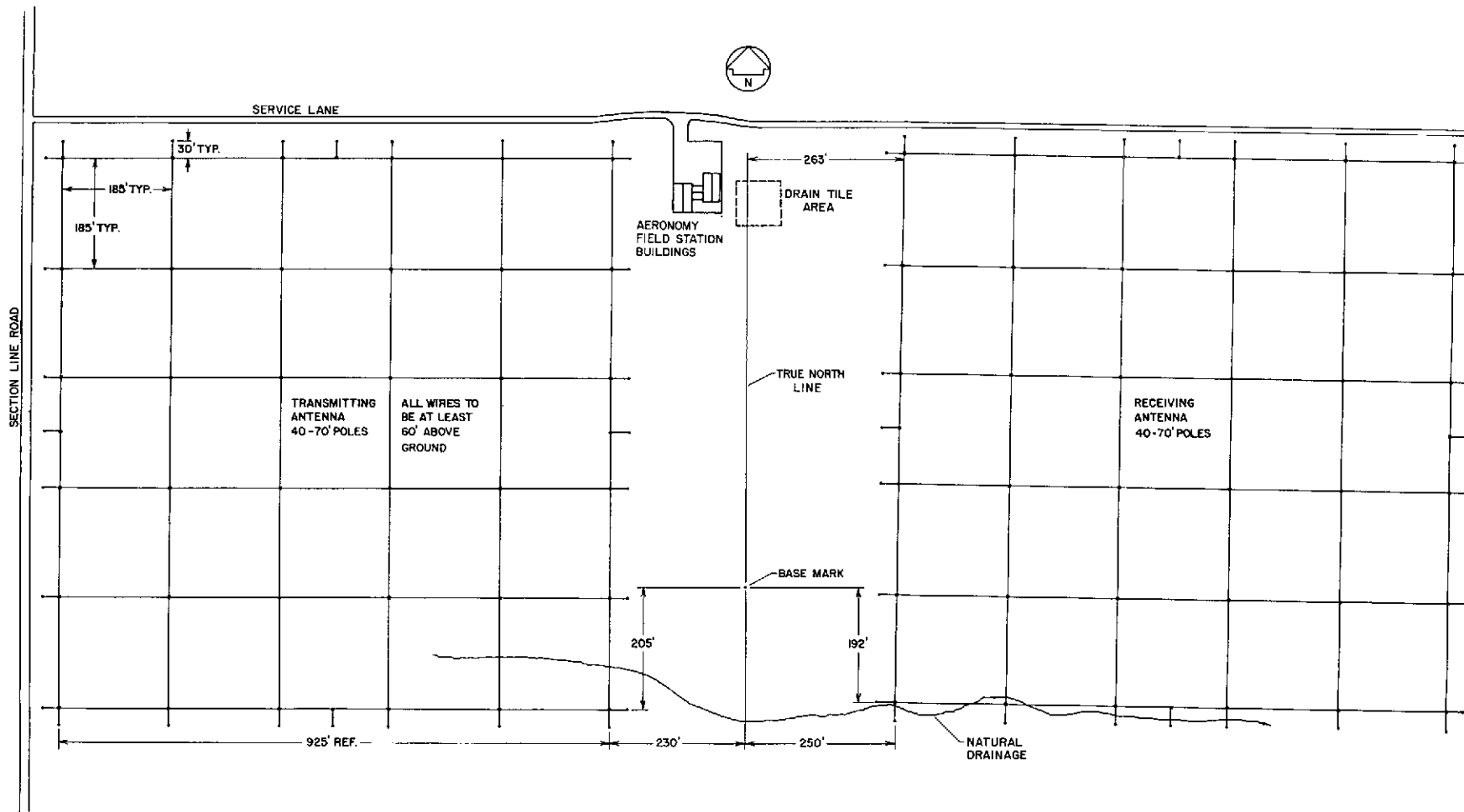


Figure 3.2 Partial-reflection antenna arrays for the Aeronomy Field Station.

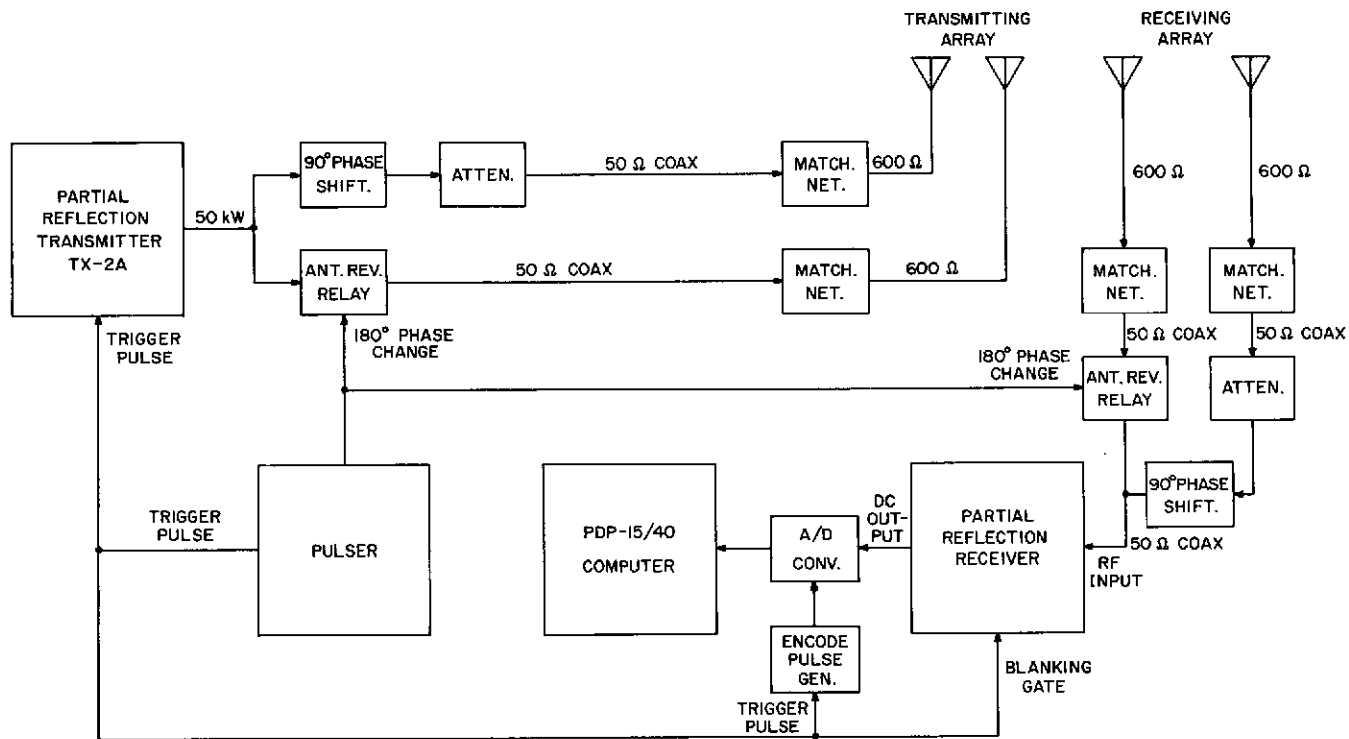


Figure 3.3 Block diagram of the partial-reflection system.

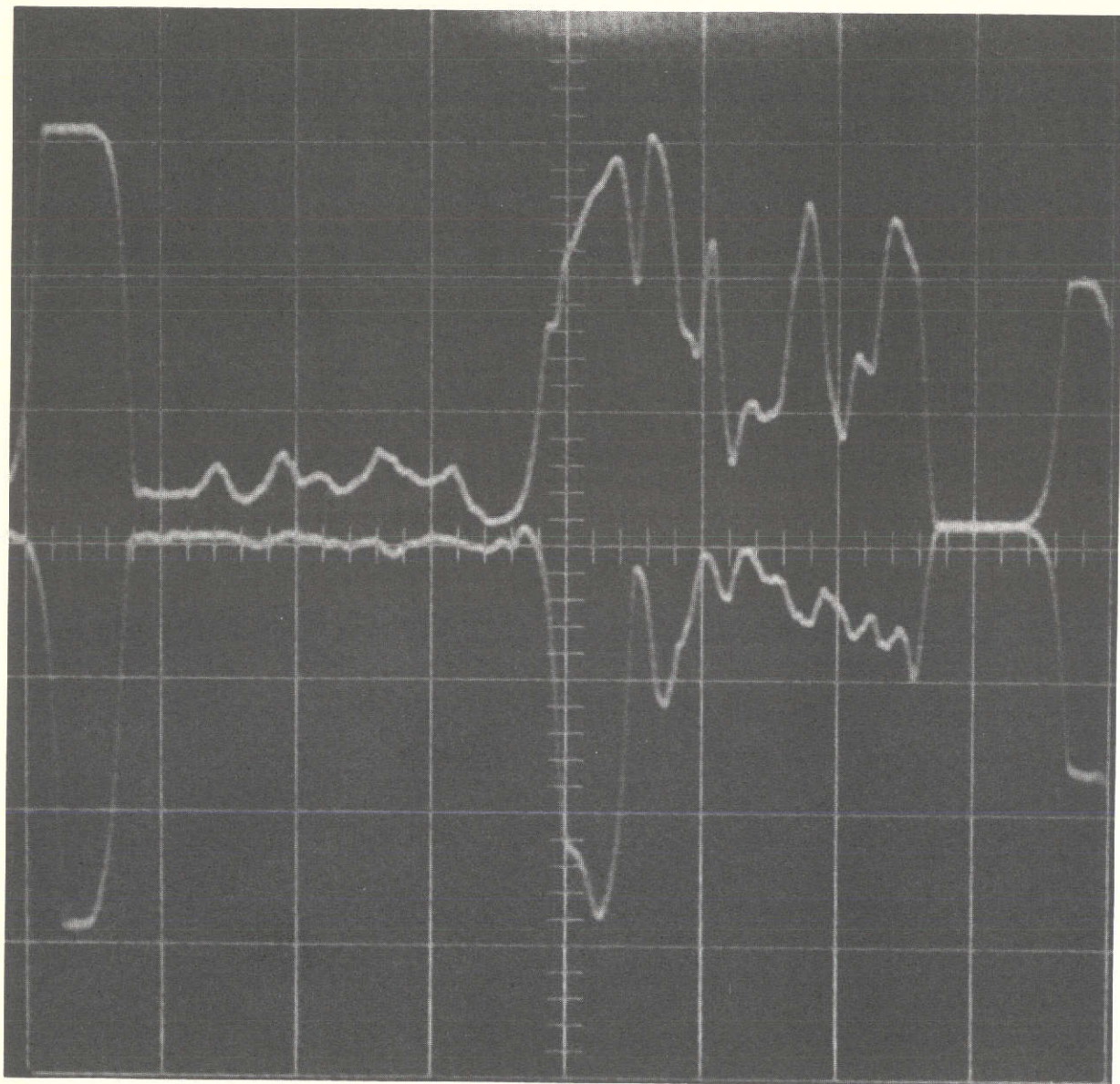


Figure 3.4 Typical frame of data as collected by *Henry* [1966].

were later measured visually and electron densities were obtained. *Pirnat and Bowhill* [1968] shows that there is good correlation between electron densities calculated from the partial-reflection data and from rocket measurements with the transmitter operating at 25 kw of power during a 50 μ sec pulse. This system of collection and storage is inexpensive, but the processing of the data to obtain electron densities is very slow and preparation and operation are complicated.

Reynolds and Sechrist [1970] set up data storage on paper tape. Ordinary and extraordinary samples were punched on paper tape for heights corresponding to 75 km and 80 km. Data can be stored at a rate of 30 values of each sample in one minute. From the paper tape the data can then be read into a computer and processed. This data on paper can be used to obtain an electron density for between 75 and 80 km. *Reynolds and Sechrist* [1970] show the results using paper tape compares favorably with results from rocket measurements and with the results published by *Belrose and Burke* [1964]. Although the system has a faster operation than the original system, it produces only one electron density and the added control circuitry is very complex.

Birley and Sechrist [1971] set up the partial-reflection experiment using a PDP-15 computer. The received signals were transmitted to the computer via an analog to digital converter and stored on DECTape to be processed later. The data consisted of four noise samples from 45 to 49.5 km and 21 data samples from 60 to 90 km in 1.5 km increments. The collection rate is 5 sets of 26 samples sec^{-1} . This collection is done alternating between ordinary partial reflection and extraordinary partial reflections. Electron densities obtained by *Birley and Sechrist* [1971] show good agreement with electron densities obtained from rocket measurements between 67.5 and 82.5 km. The other heights suffered

from too many rejections due to noise and saturation of the analog to digital converter, small signal to noise ratios, or inaccurate A_x/A_o ratios. Computer storage offers several advantages:

- 1) A fast rate of data collection (presently limited to the transmitter speed)
- 2) Data can be stored more compactly and in much larger quantities
- 3) The controlling circuitry is greatly simplified
- 4) The data processing is faster
- 5) $[e]$ can be obtained for every 1.5 km

This type of system also poses several disadvantages:

- 1) High cost
- 2) Development of computer software
- 3) Loss of accuracy in digitizing the data
- 4) Development of new circuitry and modification of the old for adaption to the A/D converter
- 5) More complicated operations (operator must know computer operation)

These disadvantages have been reduced with additional equipment and development as given in Section 3.3.

3.2 *Partial-Reflection Data Collection for the Solar Eclipse*

The partial-reflection receiver was interfaced into the PDP-15 computer to obtain data to be processed as described by *Birley and Sechrist* [1971]. Several changes in the receiver and controlling circuitry and the addition of an analog-to-digital converter were required prior to using the computer. A block diagram of the original receiver is shown on page 18 of Aeronomy Report

No. 13, [Henry, 1966]. The analog-to-digital converter saturates with an input of one volt or greater and will be damaged with inputs greater than five volts. The maximum output of the receiver was therefore reduced from 10 volts to 1.5 volts by one of the IF amplifiers, and the full-wave bridge diode detector was replaced by a single diode to reduce the nonlinearity of the receiver. A second blanking gate was inserted with the mixer in the RF amplifier module to more completely remove the initial effects of the transmitter pulse. The polarity reversal circuitry was not used but was left intact while the differential amplifier and inverter were replaced by two DC amplifiers on integrated chips.

The block diagram of the modified receiver is shown in Figure 3.5. Figure 3.6 shows the RF module with the extra blanking gate and Figure 3.7 shows the IF amplifier/DC amplifier module with the revisions. Both modules were modifications of the RF-3 module and IF-6 module respectively, given by Henry [1966]. The receiver power supply was unchanged as set up by Henry [1966]. Encode pulses as shown in Figure 3.8 were used to control the operation of the A/D converter after Birley and Sechrist [1971]. The encode pulse circuitry consists of a 5-volt power supply and 4 monostable multivibrators (Figure 3.9) with a variable timing for length of noise and signal pulses and the delay of each.

Two main modifications were made in the software set up by Birley and Sechrist [1971]. For the first change D. R. Ward [private communication] set up a computer-controlled synchronization with the external pulser. The timing shown in Figure 3.9 is used to determine which radio wave mode has been received. The computer programs are set up to store only pairs of sets of 26 numbers read from the A/D converter. A set of numbers is read in and assumed to be from a radio wave of ordinary mode. The computer's clock is set for 150 μ sec and the computer waits for another set of numbers. If another set is not read in prior

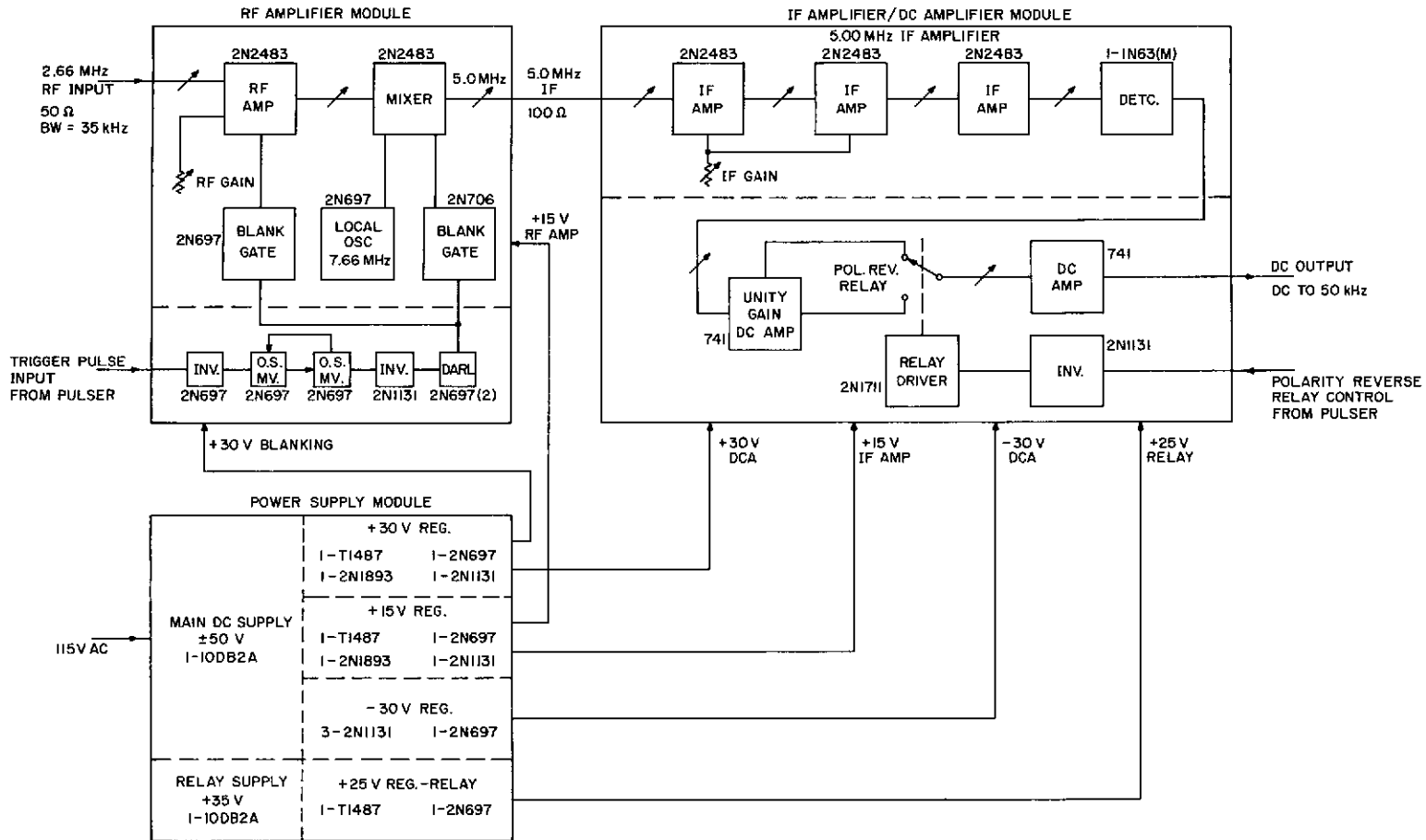


Figure 3.5 Block diagram of the revised receiver used to operate with a PDP-15 computer.

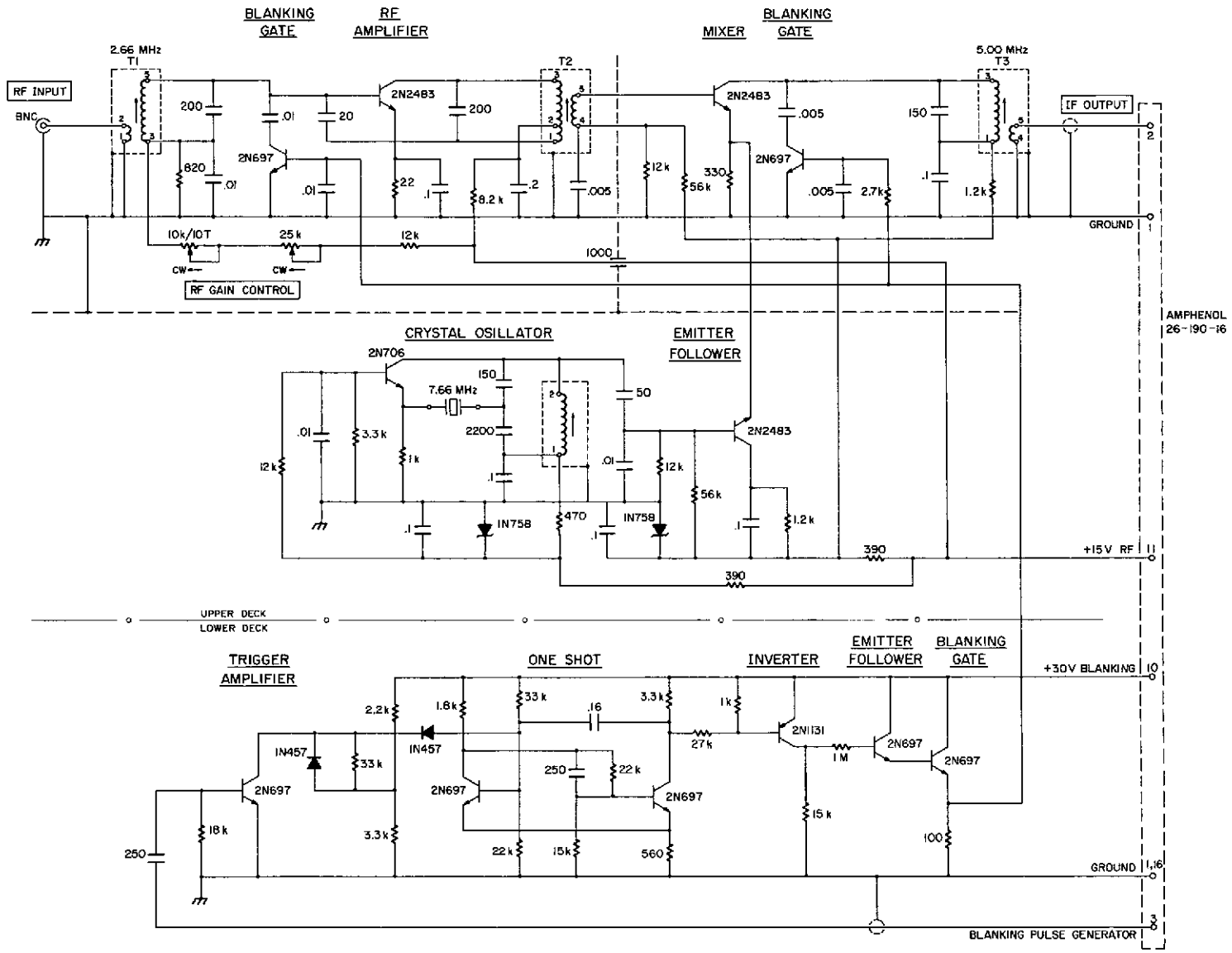


Figure 3.6 The RF amplifier module for the receiver.

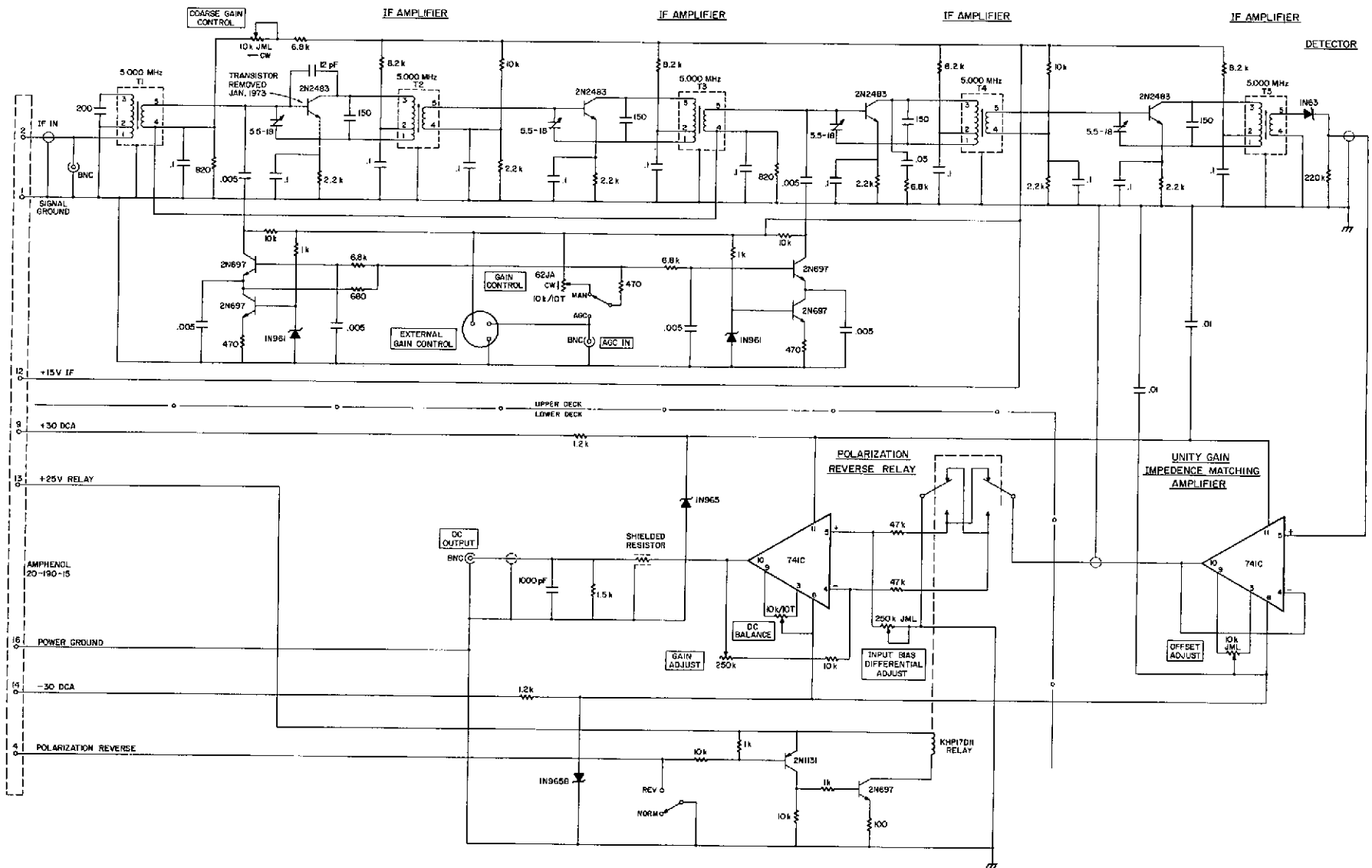


Figure 3.7 The IF and DC amplifier module.

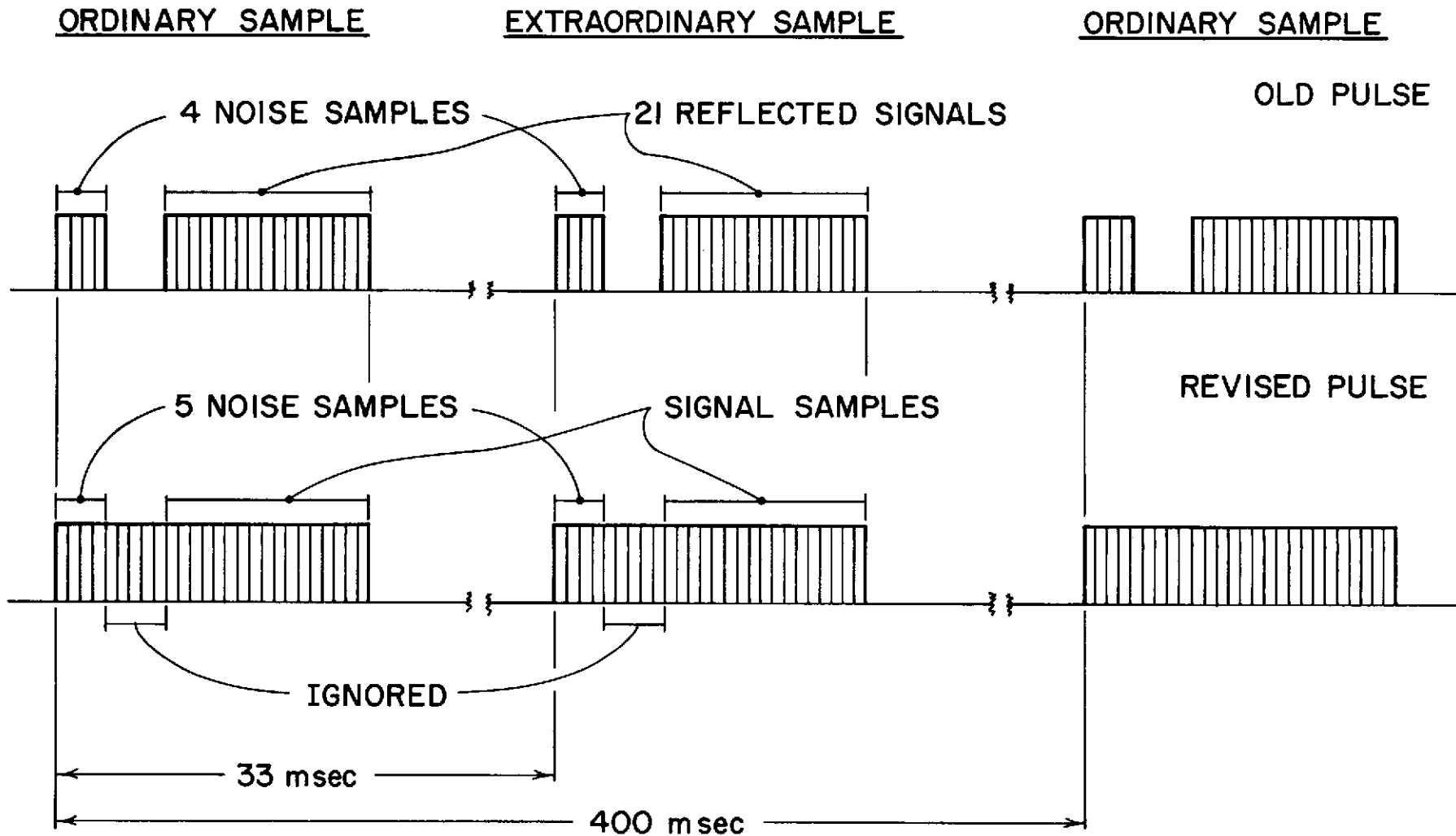


Figure 3.8 The encode pulses as set up by *Birley and Sechrist* [1971] used to collect data during the eclipse, and the revised encode pulses used by the present programs.

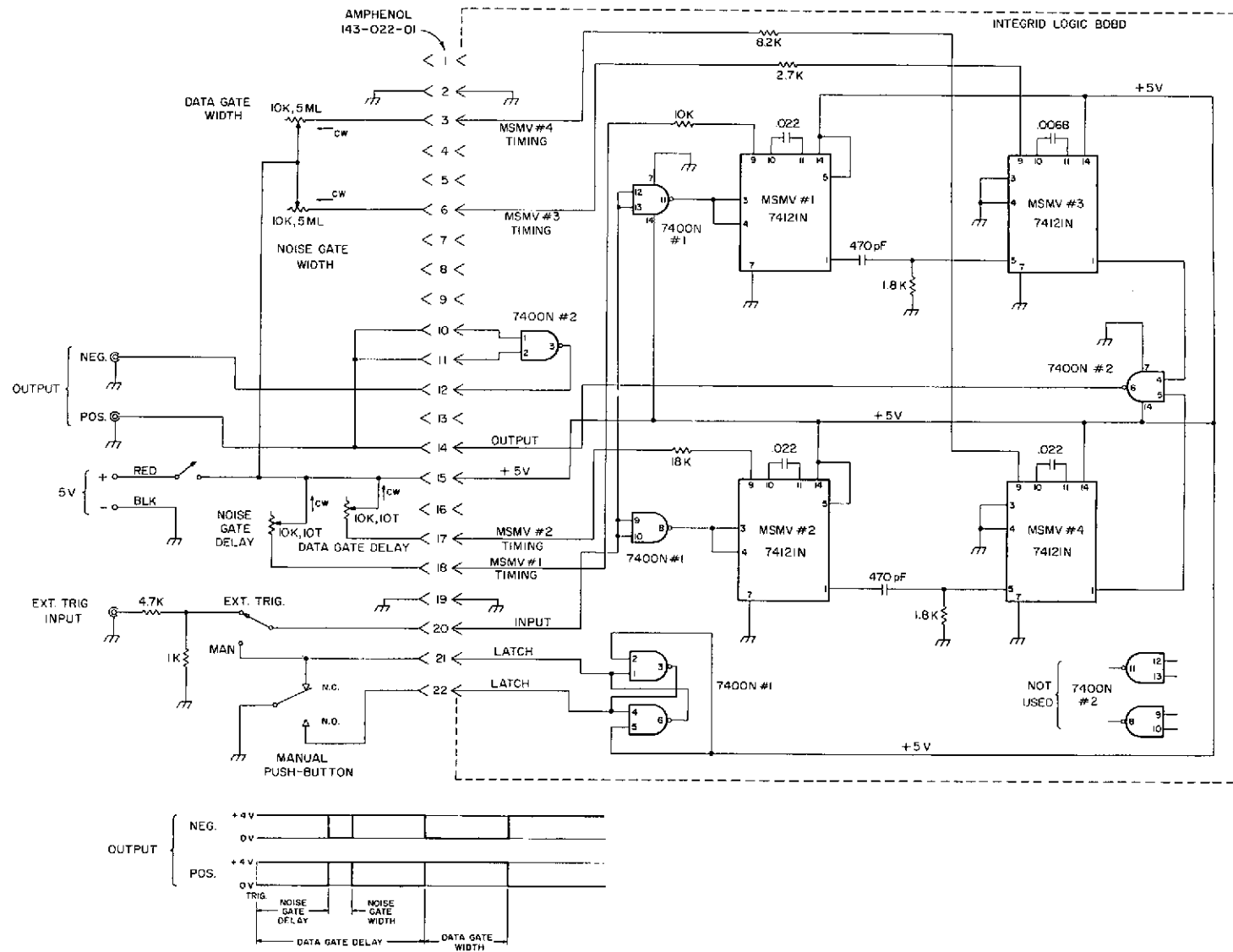


Figure 3.9 The encode pulse circuitry used to produce the former and present encode pulses.

to the 150 μ sec, the set was from an extraordinary radio wave and is rejected. Otherwise, both sets are accepted and the computer is synchronized with the pulser. This process is done only when the computer has a possibility of being out of synchronization with the pulser which are:

- 1) Beginning of every file
- 2) After the transfer of a block of data to disk
- 3) After collection is stopped and restarted by console control switch
- 4) During a timing error (no longer a terminal error, see Section 3.3)
- 5) When the computer "forgets to read" (discussed in Section 3.3)

The second change is to account for the nonlinearity of the receiver as seen in Figure 3.10 and was initially set up to adjust the data during processing [Wiersma and Sechrist, 1972]. Due to the time needed for the calibrating operation (approximately a half day), the computer is used which increases the speed of the process while making it possible to account for inaccuracies in the analog to digital converter. This process takes about 40 minutes (including 30 minutes for the receiver warm up). The adjustment to the data is done by using a table look-up method in the collection programs. Since the data stored on the disk are linearized data, the table is not needed after the collection is done and can be deleted after all the data are stored. The method is to convert the A/D converter output to the corresponding normalized receiver input. This is done by injecting a CW signal of a known value using an attenuator with one dB increments and storing the output in the computer using the set up shown in Figure 3.11. Straight line segment approximations to the curve in Figure 3.10 are obtained as shown in Table 3.1. Using outputs from 0 to 511 the corresponding inputs are determined normalized to 511 maximum, stored in a table as shown in

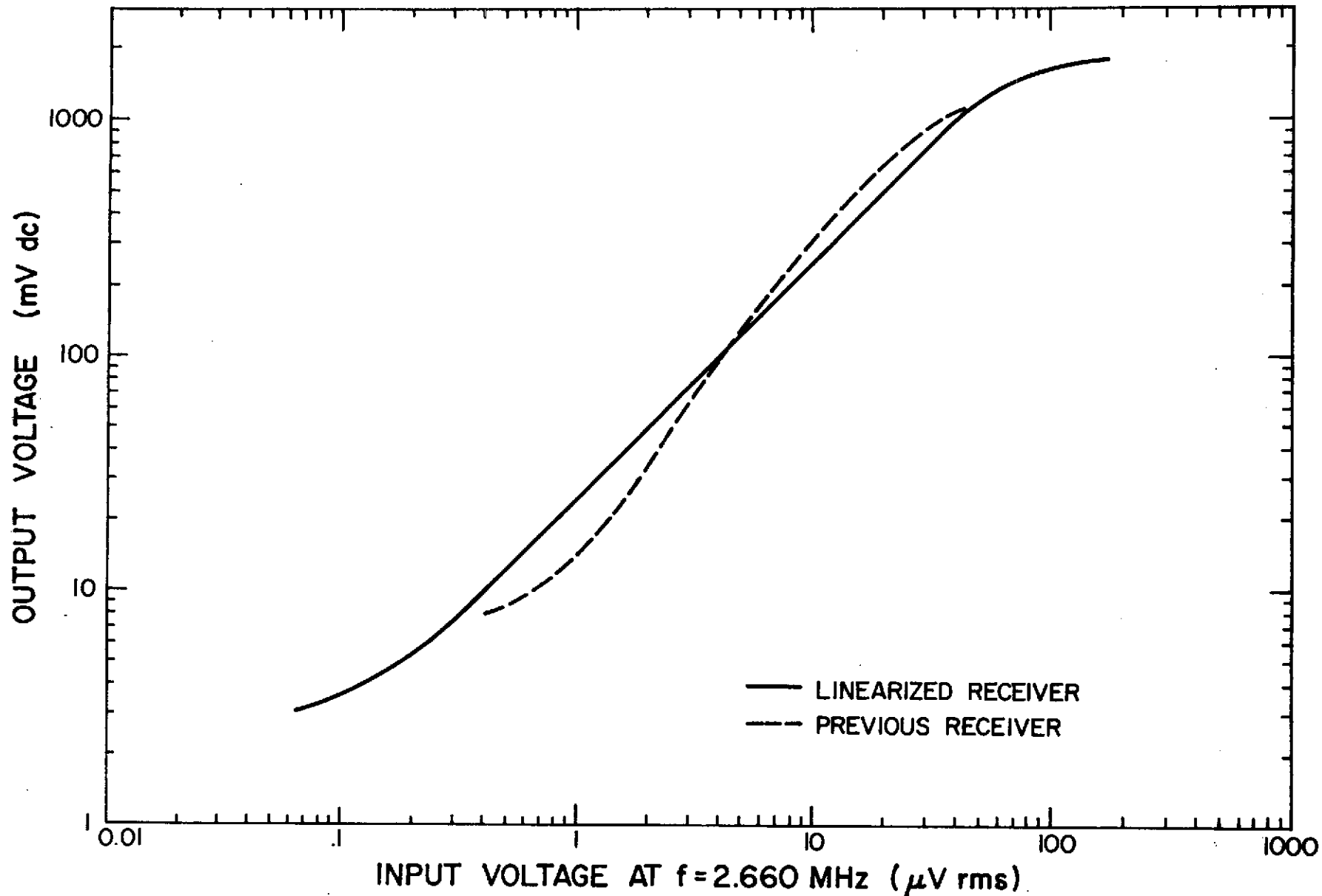


Figure 3.10 Graphs of the input versus output of the receiver used for eclipse data collection (old receiver) and the receiver presently being used. The input and output values have been normalized to the maximum of the A/D converter (511).

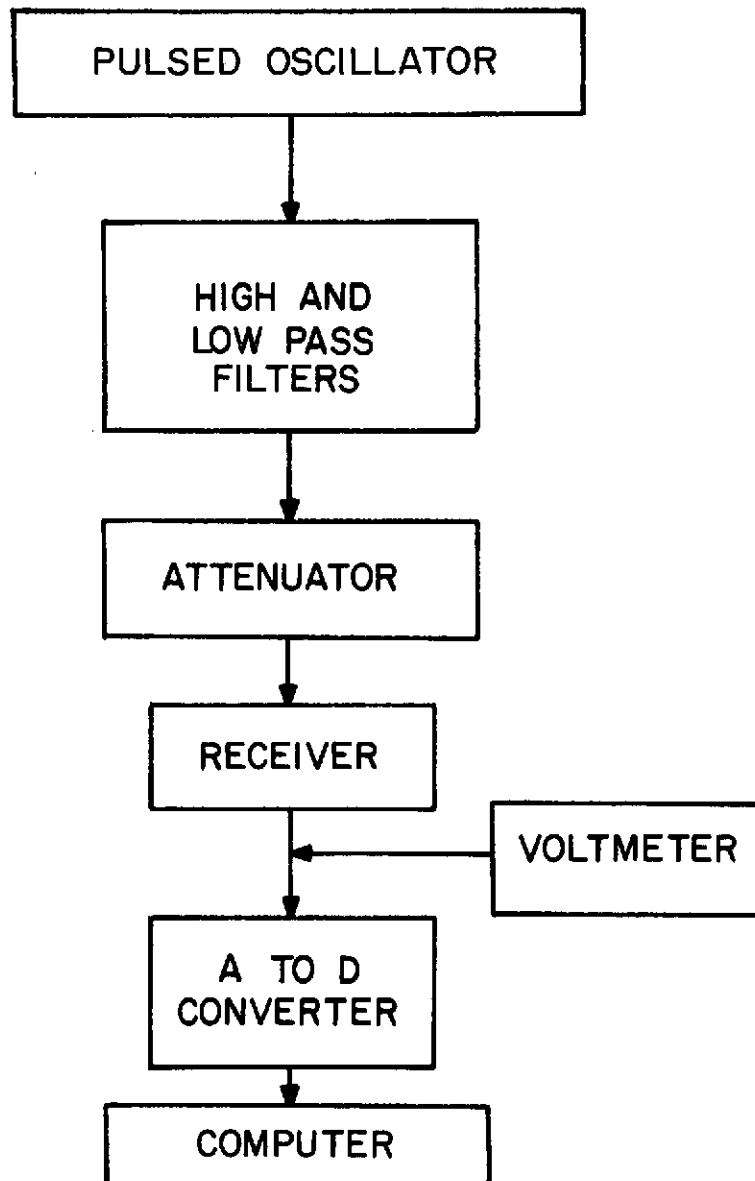


Figure 3.11 The wiring diagram used to calibrate the receiver. The voltmeter is used in setting the initial signal level prior to calibrating.

Table 3.1

Straight line segment approximation to the relationship of receiver input to receiver output.

Slope	Input	Output	Attenuation Used
S(1)=50.906	TU(1)= 0.000	TUO(1)= 4.786	99DB
S(2)= 3.912	TU(2)= 5.620	TUO(2)= 4.896	45DB
S(3)= 2.394	TU(3)= 6.310	TUO(3)= 5.073	44DB
S(4)= 7.762	TU(4)= 7.080	TUO(4)= 5.394	43DB
S(5)= 9.151	TU(5)= 7.940	TUO(5)= 5.505	42DB
S(6)= 2.576	TU(6)= 8.910	TUO(6)= 5.611	41DB
S(7)= 4.190	TU(7)= 10.000	TUO(7)= 6.034	40DB
S(8)= 1.983	TU(8)= 11.220	TUO(8)= 6.326	39DB
S(9)= 2.255	TU(9)= 12.590	TUO(9)= 7.016	38DB
S(10)= 1.676	TU(10)= 14.130	TUO(10)= 7.699	37DB
S(11)= 2.576	TU(11)= 15.850	TUO(11)= 8.726	36DB
S(12)= 1.946	TU(12)= 17.780	TUO(12)= 9.475	35DB
S(13)= 1.356	TU(13)= 19.950	TUO(13)= 10.590	34DB
S(14)= 1.754	TU(14)= 22.390	TUO(14)= 12.389	33DB
S(15)= 1.381	TU(15)= 25.120	TUO(15)= 13.945	32DB
S(16)= 1.012	TU(16)= 28.180	TUO(16)= 16.161	31DB
S(17)= 1.161	TU(17)= 31.620	TUO(17)= 19.560	30DB
S(18)= 0.969	TU(18)= 35.480	TUO(18)= 22.886	29DB
S(19)= 0.910	TU(19)= 39.810	TUO(19)= 27.353	28DB
S(20)= 1.113	TU(20)= 44.670	TUO(20)= 32.694	27DB
S(21)= 0.987	TU(21)= 50.120	TUO(21)= 37.591	26DB
S(22)= 0.771	TU(22)= 56.240	TUO(22)= 43.794	25DB
S(23)= 1.172	TU(23)= 63.100	TUO(23)= 52.693	24DB
S(24)= 0.708	TU(24)= 70.800	TUO(24)= 59.260	23DB
S(25)= 0.880	TU(25)= 79.430	TUO(25)= 71.451	22DB
S(26)= 0.832	TU(26)= 89.130	TUO(26)= 82.478	21DB
S(27)= 0.980	TU(27)= 100.000	TUO(27)= 95.540	20DB
S(28)= 0.789	TU(28)= 112.200	TUO(28)= 107.992	19DB
S(29)= 0.900	TU(29)= 125.900	TUO(29)= 125.364	18DB
S(30)= 0.816	TU(30)= 141.250	TUO(30)= 142.426	17DB
S(31)= 0.918	TU(31)= 158.490	TUO(31)= 163.559	16DB
S(32)= 0.766	TU(32)= 177.830	TUO(32)= 184.627	15DB
S(33)= 0.944	TU(33)= 199.530	TUO(33)= 212.947	14DB
S(34)= 0.913	TU(34)= 223.870	TUO(34)= 238.730	13DB
S(35)= 0.933	TU(35)= 251.190	TUO(35)= 268.645	12DB
S(36)= 1.030	TU(36)= 281.840	TUO(36)= 301.499	11DB
S(37)= 1.111	TU(37)= 316.230	TUO(37)= 334.895	10DB
S(38)= 1.260	TU(38)= 354.820	TUO(38)= 369.618	9DB
S(39)= 1.306	TU(39)= 398.110	TUO(39)= 403.979	8DB
S(40)= 1.471	TU(40)= 446.680	TUO(40)= 441.155	7DB
S(41)= 1.809	TU(41)= 501.190	TUO(41)= 478.202	6DB
	TU(42)= 562.340	TUO(42)= 512.000	5DB

Table 3.2, and placed on a storage device (normally a disk). The program DLOGF (given in the Appendix in MACRO language) reads Table 3.2 into the computer, and the table is used during collection of the received partial-reflection signal. Using the table, the MACRO subroutine LIN does the linearization of the numbers read from the analog to digital converter. The programs responsible for the formation of these two tables are TBFORL (FORTRAN IV), LINAP (FORTRAN IV), RADC (MACRO), and TTM (MACRO).

The system as it has been described was used to collect and process the partial-reflection data for the three-day eclipse period of July 9, 10, and 11, 1972. The rest of this chapter will describe further changes and developments of the system. These changes have been due to an increase of 16 K core memory, the addition of 2 disk units capable of storing 262.144 words each, and the changing from a single user monitor system to a background/foreground monitor disk system.

3.3 *Real-Time Data Storage and Automatic Processing*

A computer operates on its own timing system and if this timing system operates along with events outside the computer that affect the operation of the computer, then the computer is said to be operating in real time. For instance, if the computer reads in a set of 26 samples and is able to manipulate or process them before the next set of samples is read in, the computer is doing real-time processing; as opposed to saving the data on tape and processing it later, as done by *Reynolds and Sechrist* [1970]. With high-speed access on the disk (16 msec access time), the background/foreground system made possible real-time collection and processing of partial-reflection data. Due to the complicated timing, slow print-out, and the noise algorithm (discussed in Section 3.4), processing of the data is postponed until after the file is stored on the disk.

The background/foreground monitor system is a double monitor, multi-priority level, software system. The two monitors are separate software systems

sharing the same hardware with programs operating in the foreground system having priority. Each system has 8 automatic priority (API) levels and a mainstream level. There are four hardware levels which have highest priority. The software levels are labeled 4, 5, 6, 7, and 0 where 4 is the highest and 0 is the mainstream, the lowest. When a program initially starts running in either background or foreground, it begins on mainstream. Certain commands require a special subroutine called a real-time subroutine and is designated a priority level from 0 to 4 and stops all operation on lower priority levels (background is lower than foreground) until it exists from the level or performs an I/O operation.

With this system the partial-reflection collection and processing programs as mentioned could operate in real time, but due to several problems in the processing of data, the data could not easily be saved except in processed form. The solution used is to collect one file of data and process that file while the next file of data is being collected. After each file is collected, the operator is told what the next attenuator setting is. The collection program also checks the setting of the switches on the console to allow the operator to control parts of the collection program. Switch 0 acts as an on/off switch which causes collection to stop collecting and wait in a loop if set to 1. Switch 1 allows the background system to share the collection and processing storage device (1 disk) if the switch is set to 1. This sharing is necessary if the collected files are to be stored on DECTape. Switch 3 allows the processed data which are printed out onto the teletype to also be punched onto paper tape if the switch is set to 0. This option is presently used to allow for later plotting of the data using a programmable calculator. Switches 2 and 5 are not used at the present. The rest of the switches are used for determining the length of each file (default length is 513 pairs of sets of

26 numbers). The time of day is determined by using the clock within the computer to give the time in hours and minutes.

The flow diagram of the programs is shown in Figure 3.12. The programs are loaded into the computer and the computer's clock is set to the time of day. The operator is given the option of calibration of the receiver. The linearization table is stored on a disk and some initial information is read in. If the table read in is erroneous the operator must re-do the calibration procedures. The collection is started on priority level 6 and processing waits for the first file to be collected. After collection of the number of sets of samples set on the console switches and the operator changes the attenuator setting, the second file is collected while the first is processed and printed out. This process continues until stopped by the operator. Information used to calculate the noise threshold as described in Section 3.4 is transferred to the processing program after each file is collected and is not stored on the disk. The processing program therefore must remain faster than the collection or this information will be lost.

The processing of files involves rejecting sets of samples that are too noisy (discussed in Section 3.4), summing the squares of unsaturated data, subtracting off the sum of the squared acceptable data, and taking the square root. The resulting data are two sets of 21 samples, one of ordinary modes (A_o) and one of extraordinary mode (A_x) radio waves. This process is done in the main processing program PROC (given in the Appendix). The electron densities are calculated in CALC2 which is discussed in Section 3.5. The results are typed out on the teletype in tabular form as shown in Table 3.3.

The first line of the print-out of processed data is the heading. This gives the time the collection of the file stopped, the date, the reason for the run, and the attenuator setting for the file. The next line contains the noise threshold and the square of the multiplying constant used in the

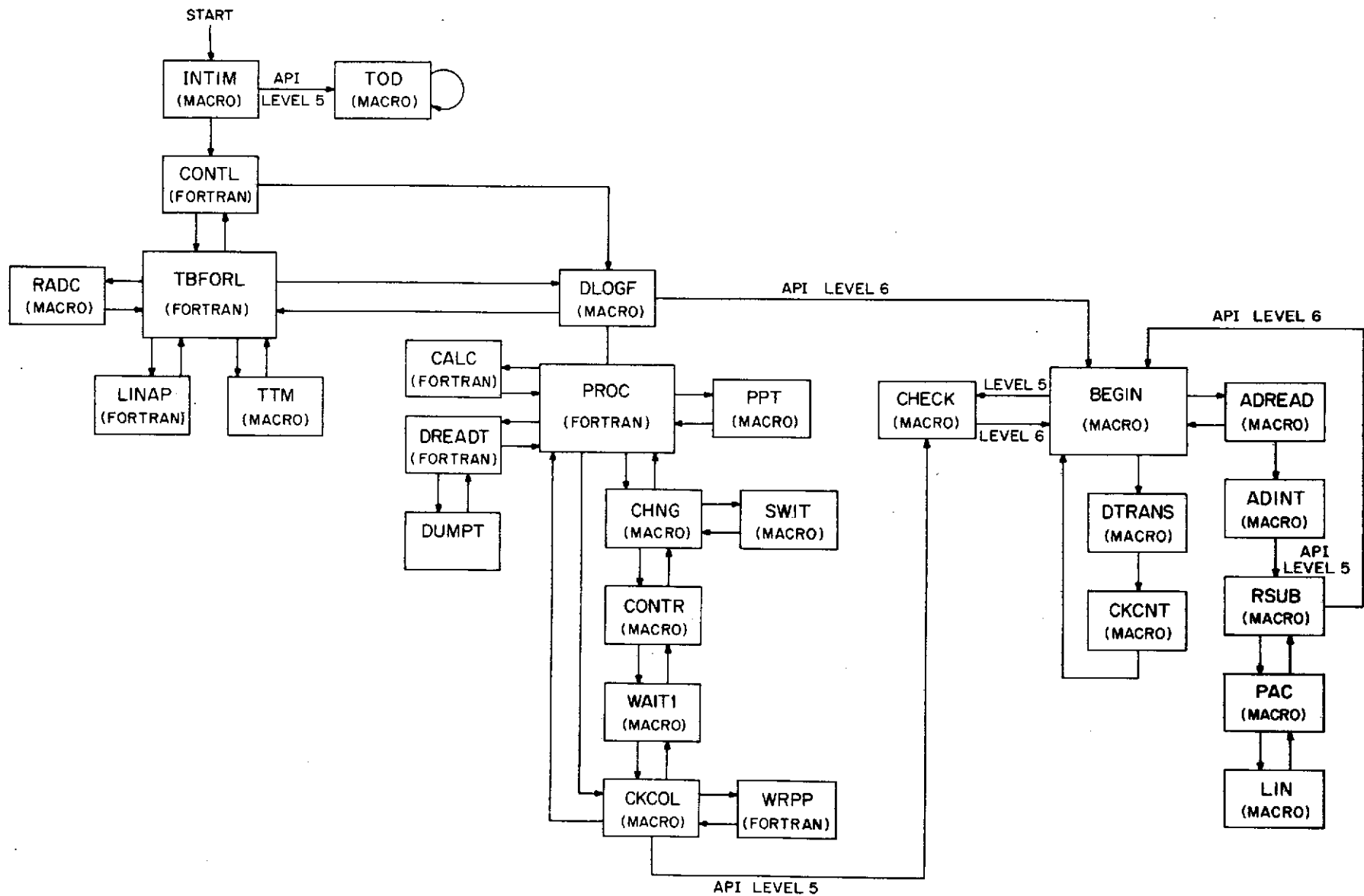


Figure 3.12 A diagram of the control flow of the partial-reflection programs. The programs operate on the API level of the preceding program unless otherwise stated. The collection and processing program operated in parallel with the collection programs operating on API levels 5 and 6, whereas everything else operates serially.

TABLE 3.3

Results of CALC2
DAILY RUN

1215

8-14-73

10DB

MAX. ALLOW. NOISE= 19.8 MULT. CONST.= 9.610

O-NOISE AV.(1) 19.9 (2) 11.6

X-NOISE AV.(1) 12.0 (2) 7.5

513 SAMPLES

44 REJ.(NOISE)

REJ. (N.+SAT.)	HEIGHT	AV. AO	AV. AX	AX/AO	ED
44	60.0	6.9	4.1	0.59	
44	61.5	9.9	2.0	0.20	5173.
44	63.0	3.3	-2.1	0.00	0.
44	64.5	3.8	4.6	1.22	0.
44	66.0	7.8	11.5	1.46	-247.
44	67.5	14.4	22.0	1.53	59.
44	69.0	22.5	34.9	1.55	105.
44	70.5	30.6	44.9	1.47	199.
44	72.0	30.8	40.0	1.30	271.
44	73.5	36.2	39.0	1.08	339.
44	75.0	54.2	49.8	0.92	298.
44	76.5	59.1	45.1	0.76	340.
44	78.0	58.9	34.1	0.58	502.
44	79.5	57.5	27.3	0.48	416.
44	81.0	56.5	20.2	0.36	682.
44	82.5	52.7	14.4	0.27	778.
44	84.0	46.6	10.8	0.23	623.
44	85.5	99.0	12.0	0.12	3003.
44	87.0	228.9	26.0	0.11	445.
53	88.5	307.2	38.0	0.12	-694.
99	90.0	324.0	41.9	0.13	-469.
					0.

maximum noise criterion discussed in Section 3.4. The next two lines are the ordinary and extraordinary mode noise before (number 1) and after (number 2) rejections due to excessive noise. The next line gives the number of pairs of sets of 26 samples collected and the number of these pairs rejected due to saturation. The first column of the table is the number of rejections due to both saturation and excessive noise for each height. The next column gives the height of the reflected signals for each row. The next two columns give RMS of the ordinary (A_o) and extraordinary (A_w) signals. The fifth column gives the ratios of extraordinary partial reflections to ordinary partial reflections from the fourth and third column respectively. The last column gives the electron density for between the heights. The last electron density is given as zero since only one height is available to calculate it.

The present method of collection and processing of partial-reflection data is fast, efficient, and easy to operate, but two problems needed to be removed. The increase of input/output operations have increased timing errors which are discussed by *Birley and Sechrist* [1971], and the A/D converter sometimes fails to respond to read commands.

The A/D converter transfers data to the computer using multicycle block transfer as described by *Birley and Sechrist* [1971]. The process is a three cycle operation for each word transferred. After each transfer, the A/D converter interface is tested for synchronization. If the timing between the interface and the I/O processor is altered, transfer is stopped resulting in a timing error. With the present system, this error can result from hardware malfunction or excessive I/O operation occurring. If the latter is the reason, the problem is only temporary and can be remedied by issuing another read. Care is taken to keep the computer synchronized with the pulser. If the error is a hardware problem, the condition will not clear up and collection must stop. The error

will usually occur when data are being collected, processed data are being printed out, and a tape is being copied onto the disk in background, all simultaneously.

The second problem has to do with the A/D converter's interface refusal to transmit data. The problem has been traced to failure in the A/D converter interface logic. The collection program will issue an A/D converter read, but not receive control back and no data are transferred. This problem occurs only with the background/foreground system and it occurs infrequently (once in about every 10,000 read commands). One solution is to issue a double read, but the problem could still occur. The solution used is for the processing to check for this stoppage, restart the collection in an orderly fashion if it has stopped and to ring the teletype bell to let the operator know of the stoppage. This solution does not prevent the failure of the A/D converter interface to transfer data, and the problem will have to be removed for faster ratio of collection, but presently the operator need not be concerned with this problem. The rest of the data is unaltered by this problem.

3.4 *Noise Rejection*

The partially reflected radio waves from the *D* region are usually small in amplitude on the order of 10 to 1000 mvolts at the output of the 80 dB gain receiver. Noise amplitudes vary between 30 to 1000 mvolts. For the purpose of the noise algorithm, noise is considered to be any interference which is part of the receiver output signal that is not attributed to the partially reflected waves from the vertically transmitted pulse. This noise is divided into two types: background noise and noise bursts. Background noise is noise caused by the receiver (14 ± 3 mV) and general atmospheric noise which is always present (40 ± 10 mV). Noise bursts are caused by lightning and other radio transmitters, and the amplitude of this noise is dependent on the location of the source. Lightning noise will usually last for the duration of one encode pulse while noise due to other

transmitters will last for at least 1/2 second which is several encode pulses (see Figure 3.8) and the noise will be increased usually by 10 to 1000 mvolts. Both types of noise are rejected in the processing program PROC(FORTRAN IV) as shown in the block diagram of this program in Figure 3.13.

Data are collected in pairs of sets of 26 numbers. Each set contains 5 noise samples and 21 samples of partially reflected signals. Each pair contains a set of ordinary mode samples and a set of extraordinary mode samples. In PROC a noise threshold is determined and the square of this multiplied by five is compared to the sum of the squares of the five noise samples of each set. This method of comparison is faster than comparing the RMS of the noise as set up by *Birley and Sechrist* [1971] since square root operations take approximately 1 msec and squaring takes 70 μ sec on the PDP-15, and the squaring need only be done once per file. If the noise of either mode is greater than the noise threshold, both sets of 21 signal samples are rejected and the next pair of sets are tested. If the noise of both modes is less than this threshold, the noise of both sets are considered acceptable and saved for later processing. The partially reflected signals with acceptable noise for each mode are checked for A/D converter saturation (.997 volts receiver output) at each height. If either of the two samples (one of each mode) is saturated at a height the two samples are rejected; otherwise the data are considered acceptable. This processing of pairs of 26 samples continues until the end of the file is reached. After the file of collected data has gone through this processing, the average of the sum of the squared acceptable noise for each mode is subtracted from the average of the sum of the squared acceptable partially reflected samples of the same mode at each height, and the square roots are printed out as shown in Table 3.3 and as described in Section 3.3.

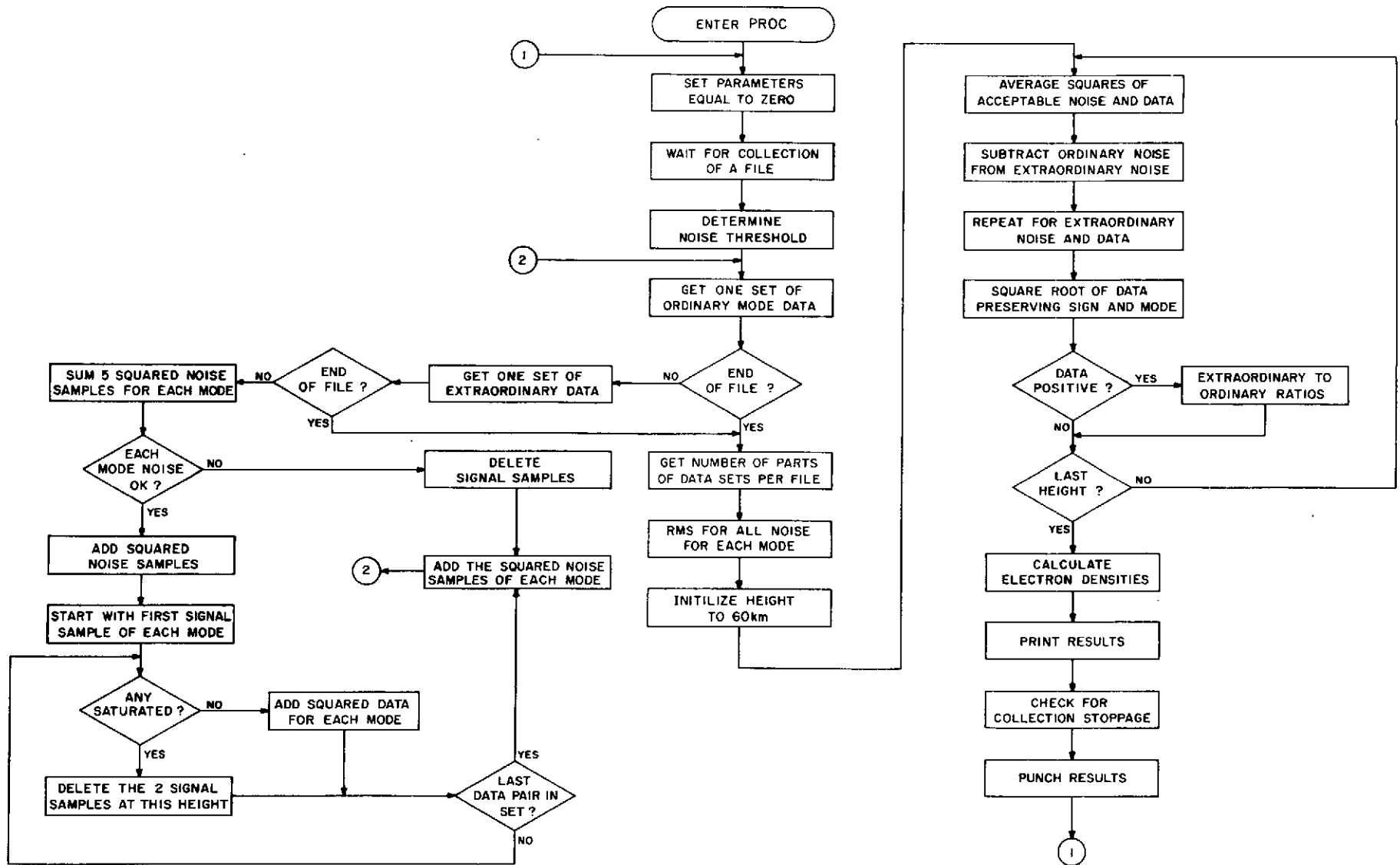


Figure 3.13 A flow chart of the processing program PROC.

Originally, the noise threshold was determined by the operator typing in a value chosen by him as seen in the program PROC73 in the Appendix. This was later changed to an automatic determination based on the attenuator setting used as given at the beginning of a run. This method did not account for the day-to-day variation in noise nor in an erroneous attenuator setting. The noise threshold value is presently determined by the following equation:

$$M = \left(K \left(\sum_0^{45} N \right) / 45 \right)^2 \quad (3.1)$$

where M = maximum allowable noise value

K = arbitrary constant

N = certain noise samples collected as explained in the following paragraph.

In the collection programs RSUB and LIN, the maximum and sum of each group of 45 noise samples are stored, and the maximum values are compared. The sum of the group with the lowest maximum value is transferred to the processing program PROC and is used in equation (3.1). The constant K has been chosen by trial and error, and values between 2.5 and 3.5 seem to give the best results (equation 3.1 is being used).

Other algorithms have been tried, but none seem to give any obvious improvement in the resulting electron densities. One method is to split 5 noise samples collected with each set of data into 2 for comparison with the noise threshold value and 3 subtracted from the reflected signals. This method works on the theory that the noise within the 5 noise samples is not the same amplitude as the noise within the 21 data samples for each set of 26 data samples, but is statistically the same over the number of samples collected

for one file. With the present system, when the number of rejections due to noise is large, (greater than 200 out of 513 pairs of sets of samples), the noise within the noise gate is restricted to a lower level than the noise in the data frame. Therefore, the noise in the data frame would not be completely subtracted off; as it would be with splitting the noise samples. The application of this technique using 4 noise samples showed no improvement in the results. Two possible causes are too few noise samples being used and the noise samples being too close together.

Another method has been developed and tested by *D. R. Ward* [private communication]. A CW signal is inputted into the receiver along with the received data from the antenna. The noise and partially reflected signals are each defined as $A \cos\theta$; where A is the amplitude and θ is the phase. The noise is assumed to be random while the partially reflected signals are assumed to have only a small variation between two sets of samples. Using an algorithm developed by *D. R. Ward* [private communication], the phase and the amplitude of the noise portion of each signal average to zero while the phase and amplitude of the signals do not. This method is used to reject the noise from the partially reflected signals at each height. This method fails to reject interference caused by other transmitted signals since this type of noise does not have random phase. *D. R. Ward* [private communication] has obtained useful electron-density profiles from the method but generally found no improvement over the present system. Further study and development of either method may improve the processing and should not be discarded.

3.5 Converting A_x/A_0 Ratios to Electron-Density Profiles

The partial-reflection programs assume a constant collision frequency for each height with seasonal variation. The values used were determined from the

following equation [Birley and Sechrist, 1971]:

$$\nu_m = Kp \quad (3.2)$$

where $K = \text{constant} = 7.3 \times 10^5$

$p = \text{pressure in pascals}$

$\nu_m = \text{collision frequency in sec}^{-1}$

The pressures used are from the mean atmospheric model from COSPAR International Reference Atmosphere (1965) with seasonal variations given by U. S. Standard Atmospheric Supplements (1966). Using these pressures, experimentally the values calculated for K vary by as much as 2×10^5 [Lodato and Mechtly, 1971]. The seasonal variations in the collision frequency (Figure 3.14) can vary by as much as 20%. This 20% variation in ν_m can cause the calculated $[e]$ to vary by a factor of 1.2. The electron densities are calculated using the refractive index equation given by Sen and Wyller [1960] and several approximations as discussed by Pirnat and Bowhill [1968]. The resulting equation given by Reynolds and Sechrist [1970] is:

$$[e] = \ln \left\{ \left[\frac{(A_x/A_o)}{(R_x/R_o)} \right]_{h_1} / \left[\frac{(A_x/A_o)}{(R_x/R_o)} \right]_{h_2} \right\} / \text{FD} \quad (3.3)$$

$$\text{FD} = (5\Delta h e^2 / 2c m \epsilon_o \nu_m) \left\{ \zeta_{5/2} \left[(\omega - \omega_L) / \nu_m \right] - \zeta_{5/2} \left[(\omega + \omega_L) / \nu_m \right] \right\} \quad (3.4)$$

where

$$\zeta_y(x) = \frac{1}{y!} \int_0^\infty \frac{\epsilon^y}{\epsilon^2 + x^2} e^{-\epsilon} d\epsilon$$

$$\epsilon = mV^2 / 2kT$$

$[e] = \text{electron density}$

$e = \text{electron charge} = 1.6 \times 10^{-19} \text{ C}$

$m = \text{electron mass} = 9.1 \times 10^{-31} \text{ kg}$

$\epsilon_o = \text{permittivity of free space} = 8.85 \times 10^{-12} \text{ F m}^{-1}$

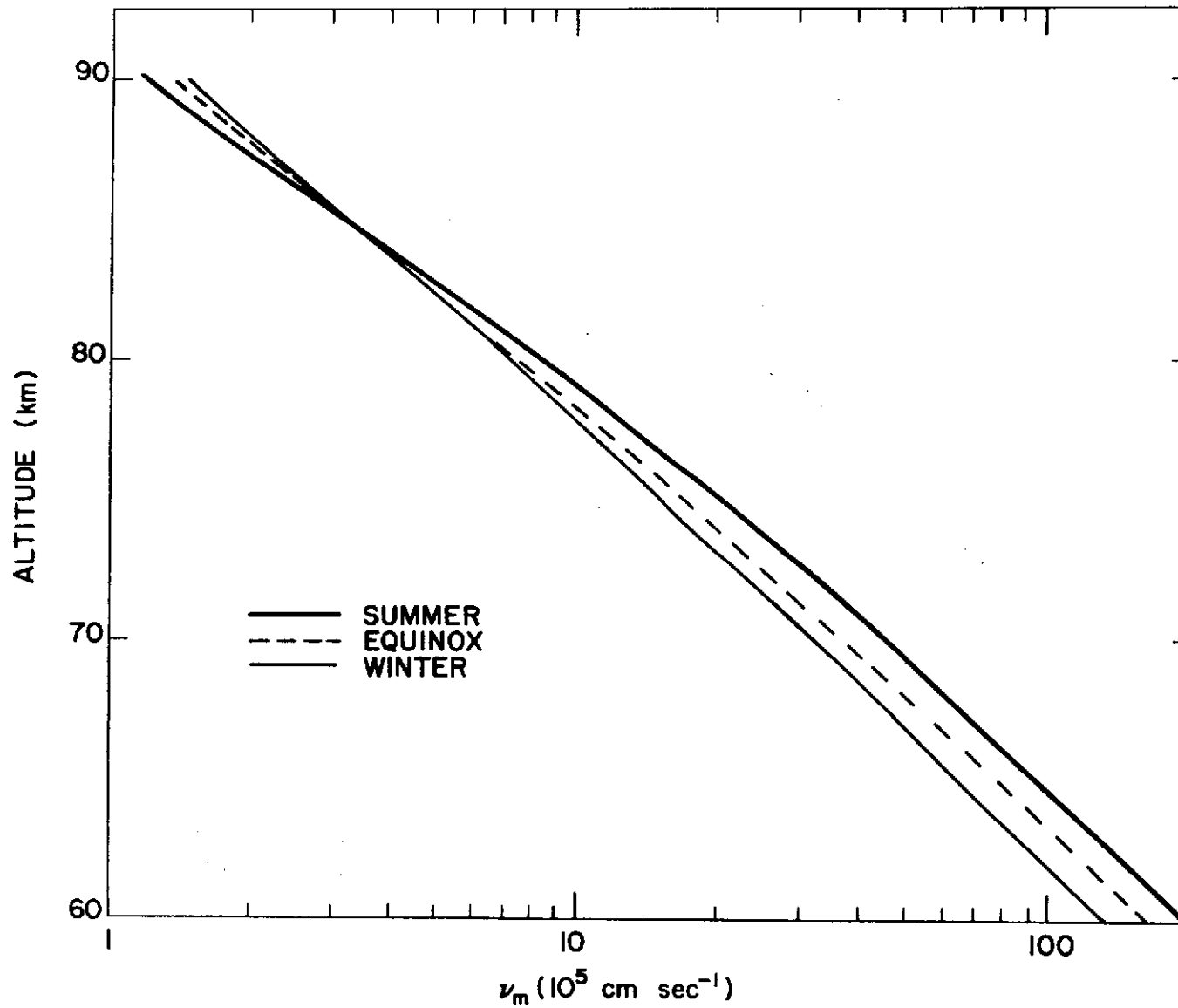


Figure 3.14 The collision frequencies used in the program CALC to obtain electron-density profiles.

- ω = angular frequency of the transmitted wave
 ω_L = gyro-frequency of the electron
 h_1 = lower height
 h_2 = higher height
 $\Delta h = h_2 - h_1$
 k = Boltzmann constant = $1.38 \times 10^{-23} \text{ J}^\circ \text{ K}^{-1}$
 T = temperature
 V = electron velocity
 R_o = ordinary mode reflection coefficient
 R_x = extraordinary mode reflection coefficient

This equation required a set of collision frequency constants which are given in the program CALC (FORTRAN IV). The ratio $(R_x/R_o)_{h_2} / (R_x/R_o)_{h_1}$ and FD (equation (3.4)) are calculated in ELDEN (FORTRAN IV). CAL2 (called by PROC) uses these values (which vary only with v_m) as constants for each pair of heights to calculate the electron densities according to Equation (3.5)

$$[e] = \ln \left(\text{RATIO2} \times (A_x/A_o)_{h_1} / (A_x/A_o)_{h_2} \right) / \text{FD} \quad (3.5)$$

where $\text{RATIO2} = (R_x/R_o)_{h_2} / (R_x/R_o)_{h_1}$.

This method is used to reduce the amount of core memory required and increase speed of execution of the program. A new CALC2 can be obtained by revising the collision frequencies and running the program CALC which writes the program CALC2. The electron densities are printed out as shown in Table 3.3 and described in Section 3.3

3.6 Equipment Testing

The equipment needs to be tested periodically to determine if it is in operating order. The transmitter is tested by observing and keeping a log of

the voltage and current at various locations via meters and an oscilloscope. The antennas are tested by transmitting and receiving signals at various times during the day. At noon the extraordinary signal should be absorbed and at night the ordinary signal should be absorbed. By transmitting and receiving ordinary and extraordinary signals as described in Progress Report 73-1 [Edwards, 1973], the phase and attenuation of each antenna of each array can be set and checked for possible damage. This process is also a partial check for the transmitter and receiver. A spot check of 30 dB difference in ordinary and extraordinary reflections from the E region at noon is done on a daily basis.

The program CHECK (FORTRAN IV) has proved valuable in checking the receiver and the analog to digital converter. CHECK performs a modified dump of the A/D converter as read by the computer. If the number 31 is typed, the output is in the form of partial-reflection data (ordinary and extraordinary pairs), patterned after the new encode pulse shown in Figure 3.8. If any other number is typed in an average of that number rounded to the next higher multiple of 50 is printed out. The 31 pairs of samples are printed out in millivolts only, while the averages are printed out in millivolts and as represented in the A/D converter. This program has had many applications; it showed the blanking gate on a new receiver to be too long. It was used to calibrate the A/D converter using an input from a standard source. Table 3.4 shows the accuracy of the A/D converter as the standard voltage source was varied from 1.0 volts to .1 in .1, .01, and .001 volt increments. It was used in comparing the paper punch system set up by Reynolds and Sechrist [1970] with the computer storing method presented in Section 3.3. CHECK has also been used to determine the number of samples required to have less than 10% error due to noise (at least 100 samples are required). The program is easy to operate and has become important in testing and checking the receiver and the analog to digital converter.

Table 3.4
The output of the A/D converter using a
calibrated input source

<u>Average</u>	ADC <u>Voltage</u>	Input <u>Voltage</u>
511.204	998.444 mV	1000 mV
460.558	899.527 mV	900 mV
409.625	800.050 mV	800 mV
358.528	700.250 mV	700 mV
307.057	599.720 mV	600 mV
256.020	500.040 mV	500 mV
205.252	400.883 mV	400 mV
154.082	300.941 mV	300 mV
102.787	200.756 mV	200 mV
51.076	99.758 mV	100 mV
46.349	90.525 mV	90 mV
40.843	79.772 mV	80 mV
35.208	68.766 mV	70 mV
30.844	60.242 mV	60 mV
25.769	50.330 mV	50 mV
19.976	39.016 mV	40 mV
15.022	29.339 mV	30 mV
10.200	19.922 mV	20 mV
4.830	9.433 mV	10 mV
3.857	7.533 mV	9 mV
3.233	6.315 mV	8 mV
3.010	5.880 mV	7 mV
2.847	5.560 mV	6 mV
2.443	4.771 mV	5 mV
2.054	4.012 mV	4 mV
1.404	2.743 mV	3 mV
0.659	1.286 mV	2 mV
0.125	0.244 mV	1 mV
0.010	0.020 mV	0 mV

3.7 Future Development

Several improvements are being made to the system. A new receiver is being made using a linear detector and new RF and IF stages to reduce the receiver noise. Figure 3.10 shows a comparison of the input versus output between the new receiver and the old one. With no input signal, the noise level of the new receiver is 2.5 mV and the level of the older receiver is 14 mV. The circuitry and discussion of it are given in the Aeronomy Progress Report 73-1 [Edwards, 1973].

A digital input/output device is presently being sought which would improve the calibration time and free the operator for other tasks as well as simplify the operation of the system. The purchase of such a device would also reduce the amount of paper presently required.

Another asset would be a line printer. One could reduce the processing time by at least half and allow for more sophisticated processing (with possibly better noise rejection) if such a line printer were purchased.

As mentioned by *Birley and Sechrist* [1971], an increase of transmitter power is also needed. This would improve the signal-to-noise ratio and give better data below 70 km.

The noise problem should be studied more carefully. Perhaps a combination of the method discussed in Section 3.4 would improve the results. Another possibility would be to reject extremely low values of reflected signals.

An additional program to transfer collected data to tape would be helpful. The original programs set up by *Birley and Sechrist* [1971] saved data on tape for future processing. With the present system, collected data can be stored on tape by using a system program called PIP. This requires knowledge in operation of the computer, and the transferring of files can get complicated.

4. EXPERIMENTAL RESULTS

This chapter describes the results from partial-reflection data which was collected and processed by the computer on July 9, 10, and 11, 1972. A solar eclipse occurred on July 10, 1972. The obscuration function shown in Figure 1.1 shows the first contact to be at 1319 CST and the last contact to be at 1536 CST with 60% of the solar disk obscured. The data were collected from 1200 to 1700 CST to show the effects of the solar eclipse on the electron density and collected between the same times on July 9 and 11 to be used as control data. Data were collected in blocks called files. Each file of data, consisting of 1026 sets of 26 numbers, was collected and stored on DECTape every 3.8 minutes. The signal prior to entering the receiver was attenuated with four attenuator settings (0, 10, 20, and 30 dB). Each file was collected beginning with the lowest attenuator setting of 0 dB with each subsequent file collected at the next attenuator settings; 10, 20, and 30 dB, respectively. This process was then repeated. This process was used to obtain the very small echoes as well as the very large ones. The files of data are divided into approximately 15 minute intervals, corresponding to the four attenuator settings.

The data between 1400 and 1430 on July 9 was lost due to an erasure of the disk before it could be processed. These data have been interpolated. The data from July 10 between 1200 and 1300 was erroneous and therefore has been eliminated from the results. The computer results were processed further combining the files with different attenuator settings.

4.1 *Reduction of Data*

Individual results shown in Figure 4.1 show valid electron densities but are limited height range; therefore, multiple attenuator settings were used to obtain usable data over a greater range of heights. The computer processes

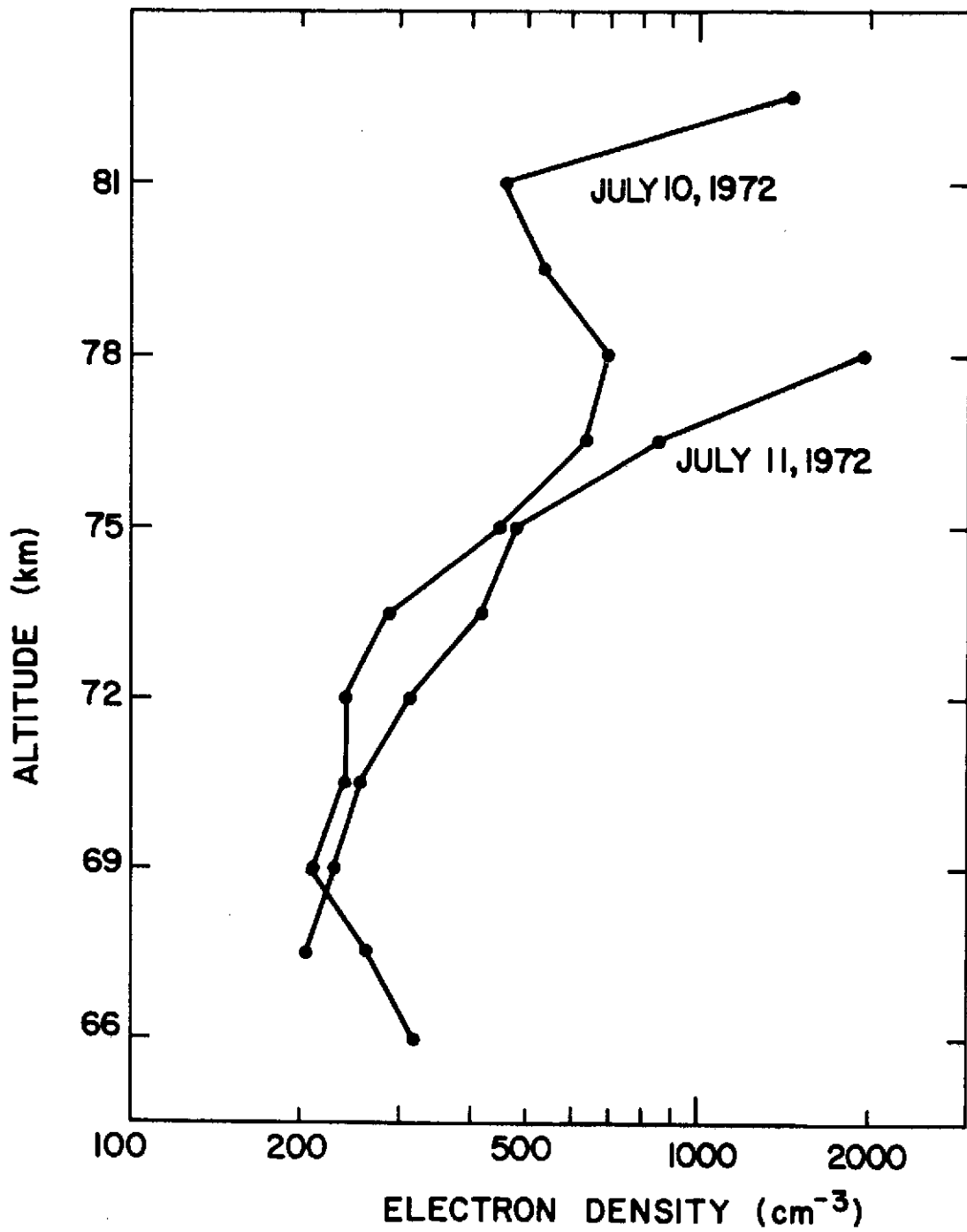


Figure 4.1 Comparison of electron-density profiles on July 10 and 11, 1972. The data were taken at 1432 CST with the attenuator set at 30 dB.

only one file at a time; therefore, further processing was necessary to combine four files corresponding to the four attenuator settings into one set of results. Three methods have been developed to accomplish this. The first method was originally used but problems developed in determining acceptable data and method two was used. Using method two, some acceptable data were being ignored and the 20 and 30 dB settings were found to give similar results. Therefore, method three was developed to utilize much of this acceptable data that were being ignored.

1. In method one, the results with the lowest attenuator setting (0 dB) were used for 60 km up to the height where 5% of the ordinary and extraordinary data was rejected due to saturation (see Section 3.4). The electron densities for the higher heights were obtained from the next higher attenuator setting under the same restrictions of saturations. This process continued until the last electron density was obtained. The results of this method seemed to be satisfactory except for above 81 km and below 66 km.
2. Method two is the same as method one, but accounts for inaccuracies in the receiver by rejecting electron densities that used A_x/A_0 ratios that were less than .09. Electron densities were rejected also if the signal to noise ratio was less than 1. These two revisions eliminated much of the results below 65 km and above 85 km.
3. Method three is similar to method two except for the way the multiple attenuators are combined. The electron densities are considered acceptable if the A_x/A_0 ratios for both heights are greater than .08, the signal to noise ratio is above 1 for both heights, and the rejections due to saturations were less than 5% for both heights used to calculate the

electron density. If more than one attenuator setting had acceptable electron densities for between two heights, then the median of the acceptable electron densities was used. Using these three methods, the computer results were combined to give one electron-density profile for every 4 attenuator settings. Using either average or medians, electron densities of different heights or of different times were combined as discussed in Section 4.2.

4.2 *Electron-Density Results*

The results are presented in two forms: by the total differential absorption below each height (A_x/A_0 ratios) and by electron densities. The A_x/A_0 ratios given in Figures 4.2, 4.3, and 4.4 are plotted using a sixth order polynomial approximation of the ratio as calculated by method one. The eclipse shows a reduction in absorption which indicated a reduction in electron density as expected. The third day shows irregular absorption with a large increase in absorption. Referring to Figure 2.2, the increase in absorption is related to the X-ray flux burst. The electron density for above 75 km for the three days given in Figure 4.5 shows a good correlation between the large increase in electron density on July 11 and the burst of X-ray flux. Due to this obvious contamination, the second control day is not used for comparison during the burst period.

Figure 4.6 gives the A_x/A_0 ratios versus height. The ratios were determined using method one and taking the median of the groups within the hour corresponding to the maximum obscuration of the solar eclipse (1400-1500 CST). Due to the much larger absorption in the control days than during the eclipse, the electron densities above 81 km (approximately) are not valid according to method two and three, but with the eclipse day, the values should be acceptable up to 85 km.

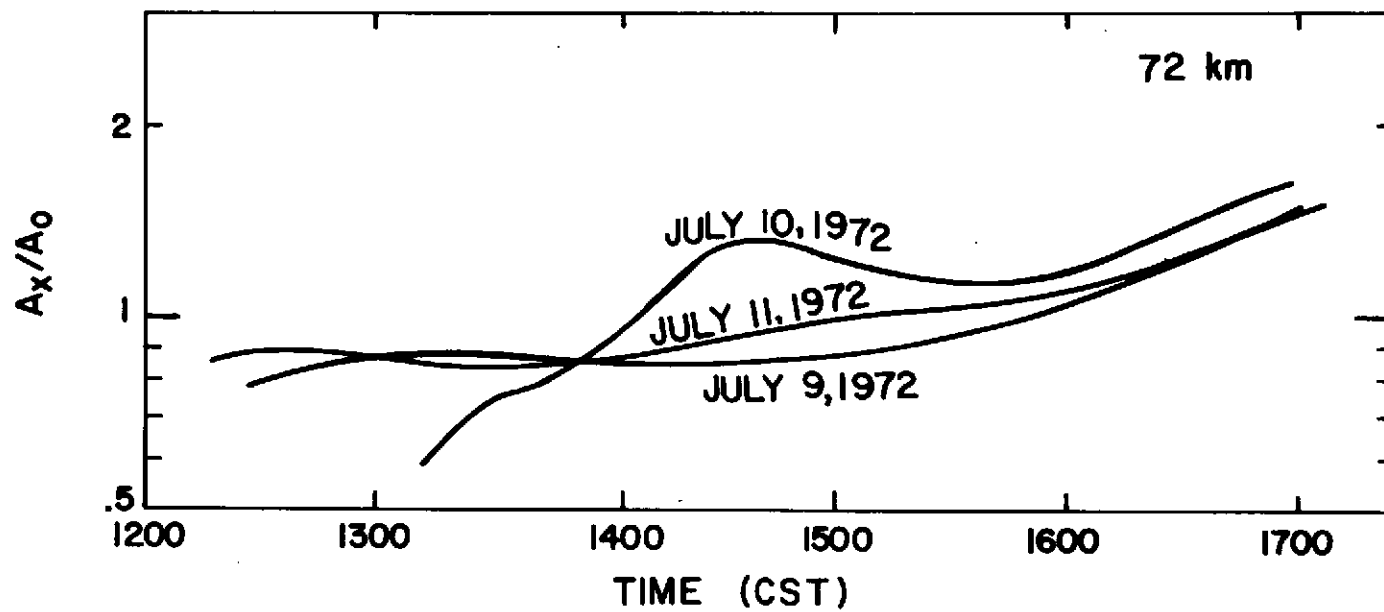


Figure 4.2 Comparison of the A_x/A_0 ratio at 72 km for July 9, 10, and 11.

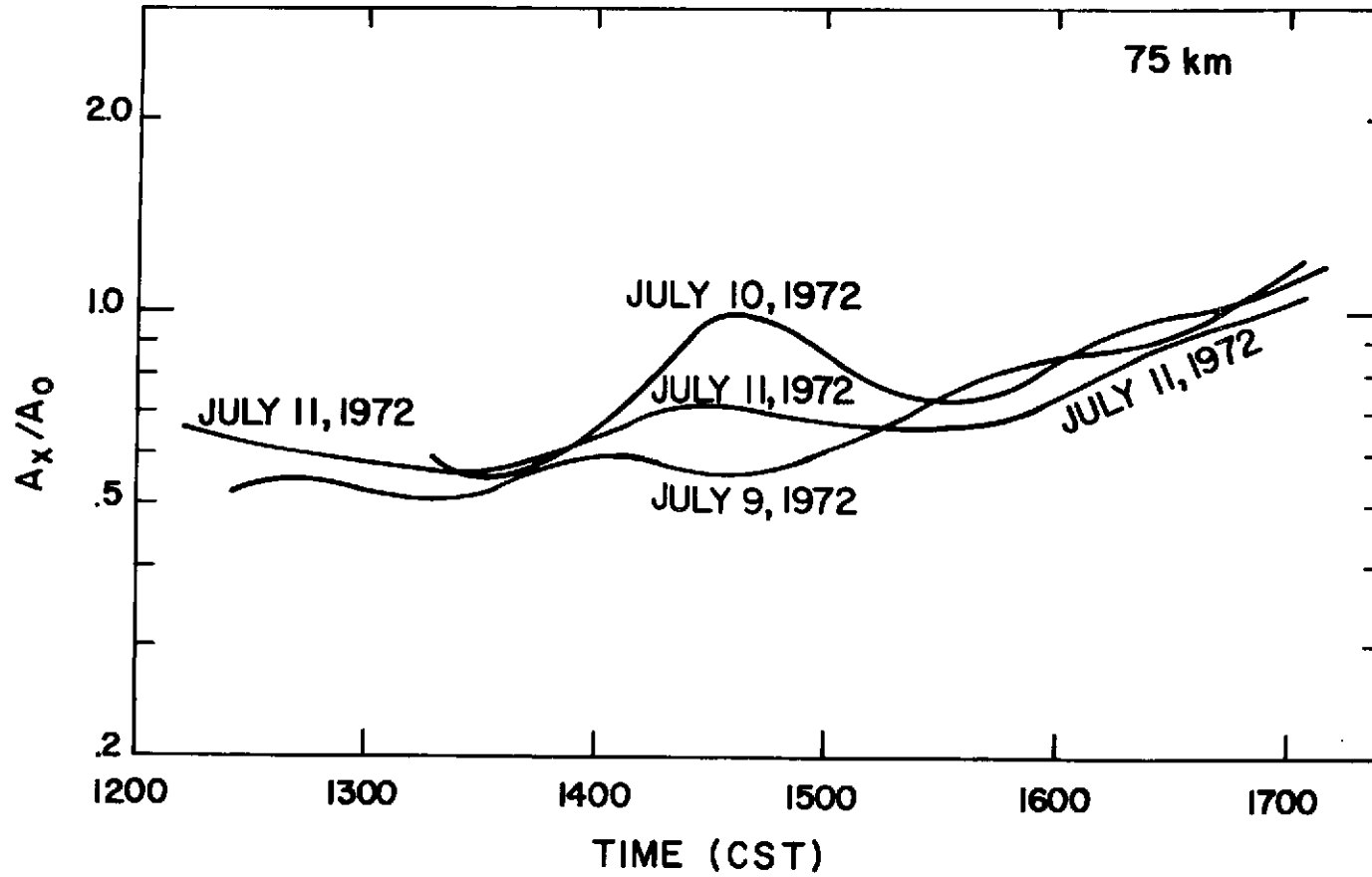


Figure 4.3 Comparison of the A_x/A_0 ratio at 75 km for July 9, 10, and 11.

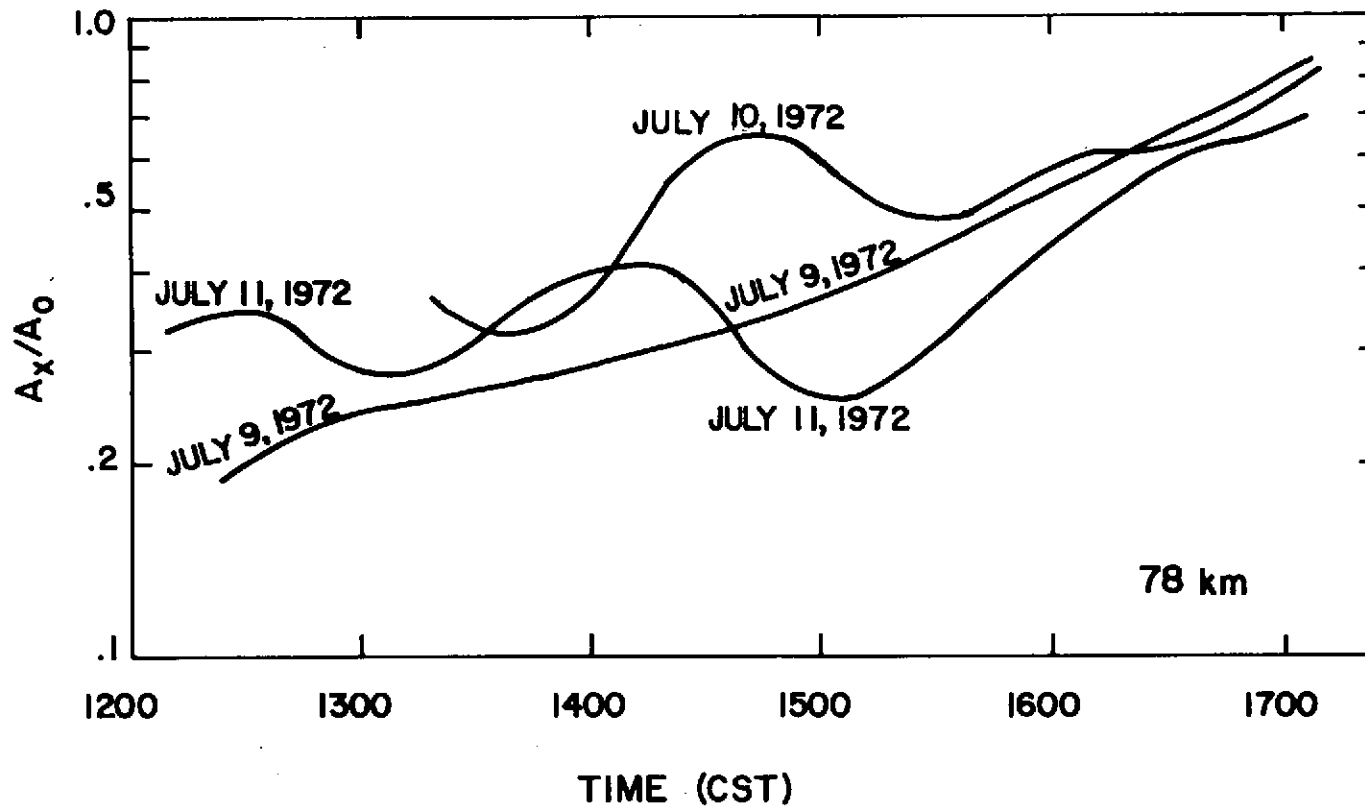


Figure 4.4 Comparison of the A_x/A_0 ratio at 78 km for July 9, 10, and 11.

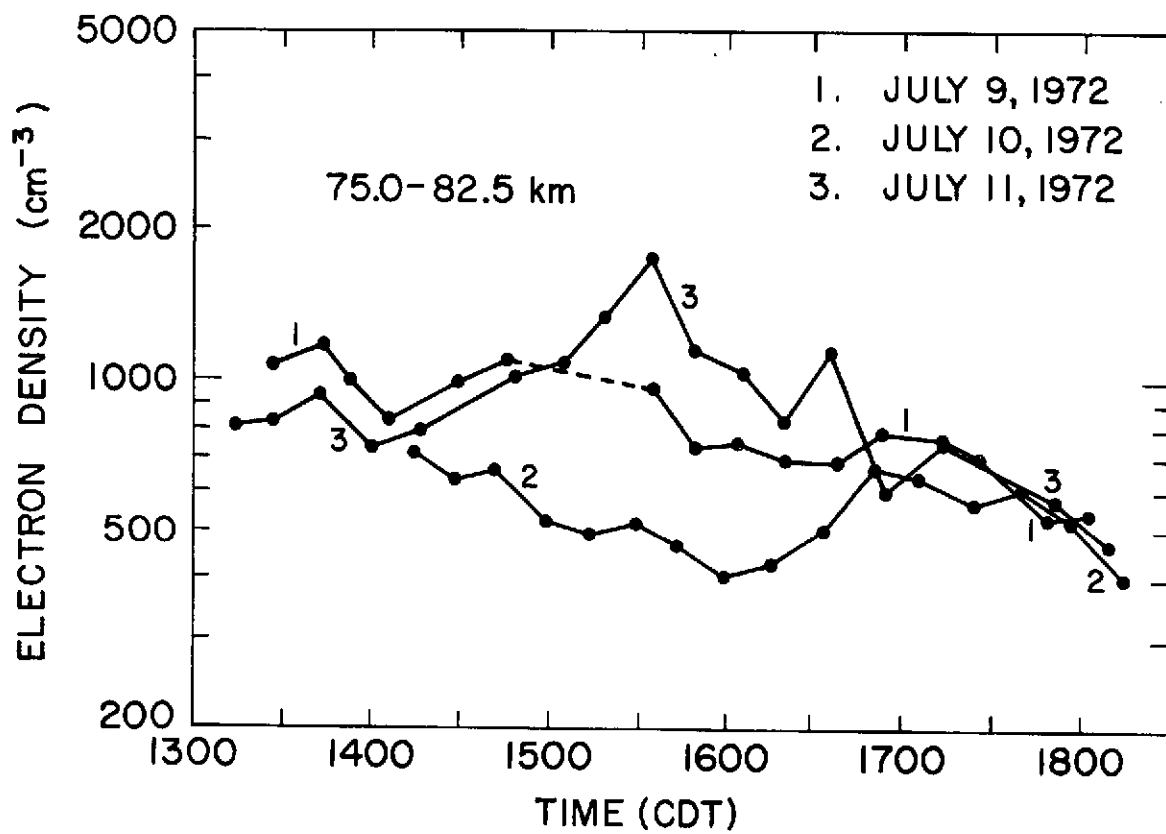


Figure 4.5 Median electron densities between 75 and 82.5 km.

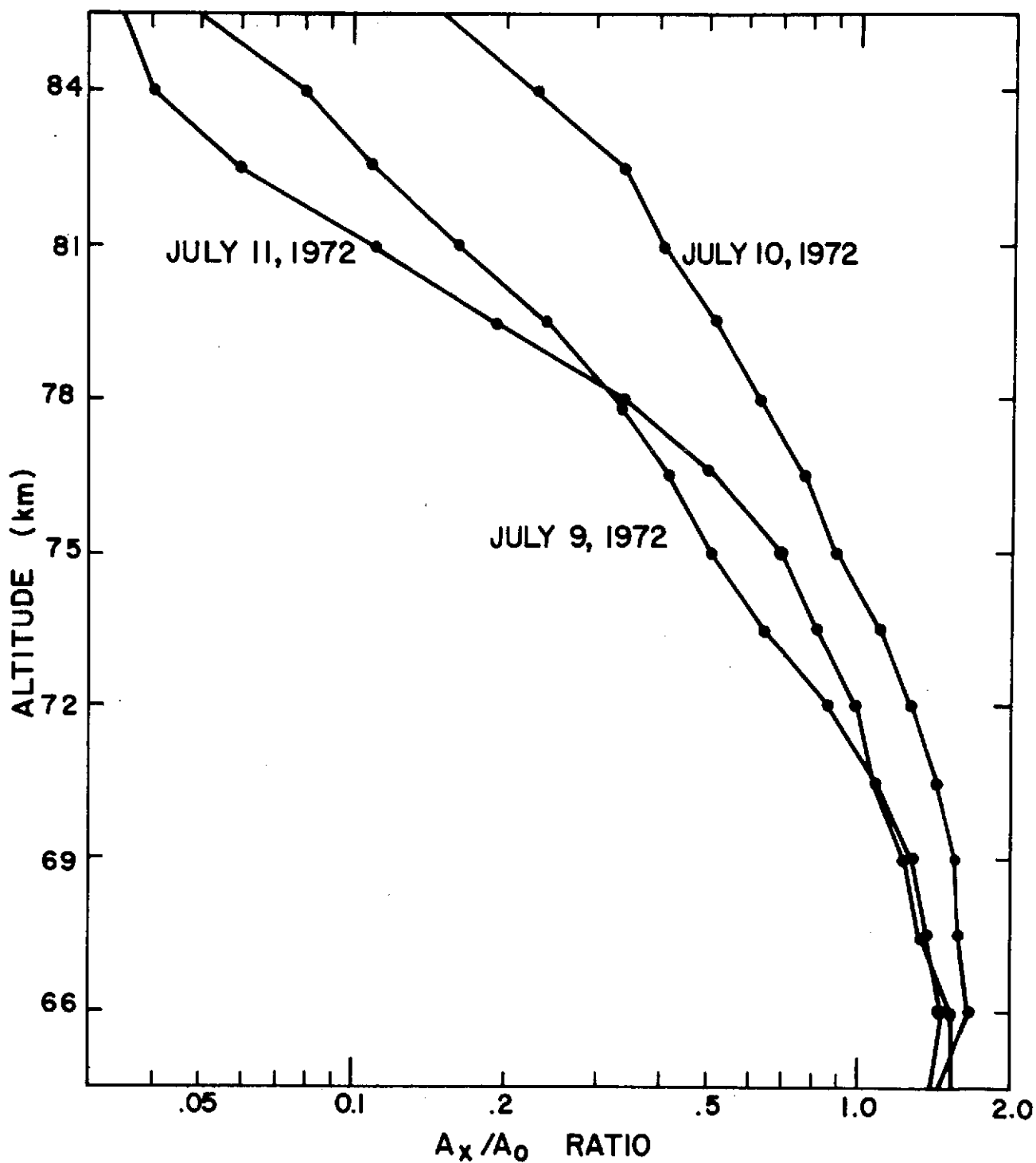


Figure 4.6 Median A_x/A_0 profiles between 1400 and 1500 CST for each day.

Figure 4.7 gives the electron density variation with time. The electron densities are averages between 70.5 and 78 km and between 78 and 87.5 km with the electron densities obtained by using method one for processing the computer result. Figure 4.5 and 4.8 give the electron-density median for 75 to 82.5 km and 67.5 and 75 km, respectively, as each varies with time. These electron densities were obtained using method three. At the lower altitudes, the median electron densities show no effect from the eclipse while the average electron densities do show a slight effect. This difference, though, is mainly attributed to the higher heights the averages were taken from rather than to the method used. The highest heights show large effects due to the eclipse. Figure 4.7 shows a minimum electron density near maximum obscuration of the eclipse while Figure 4.5 shows the minimum being delayed by half an hour. This is attributed to the variation in the data due to the inaccuracies in the partial-reflection equipment. The X-ray burst shown in Figure 2.2 seems to have no effect at the lower altitudes.

Median electron-density variations with height are given in Figure 4.9. These values are the median obtained by processing the computer results utilizing method two and finding the median value between 1400 and 1500 CST. Below 75 km the eclipse does not seem to have much effect on the electron density as shown in Figure 4.9, but above 75 km, the electron density decreases by 45 to 65%. The upper height for this comparison is 81 km due to the small A_x/A_o ratios (shown in Figure 4.6). The electron-density profile shows some conformity to the expectation given in Section 2.4.

4.3 *Theoretical Applications*

Since the eclipse never reached totality, the electron production (q) cannot be assumed to be zero, but equation (2.4) can be used as an approximation

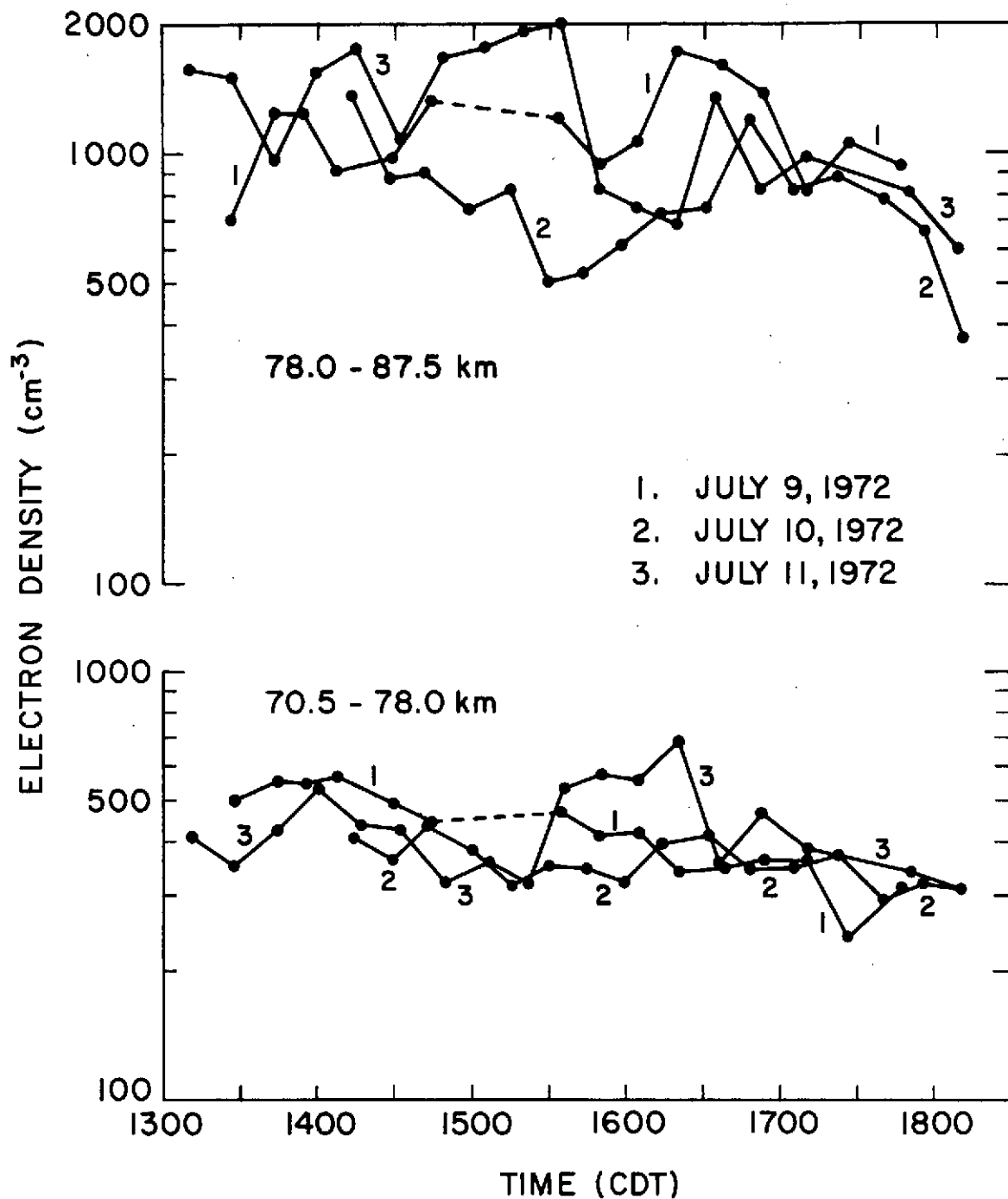


Figure 4.7 Average electron densities between the altitudes 78.0 - 82.5 km and 70.5 - 78.0 km.

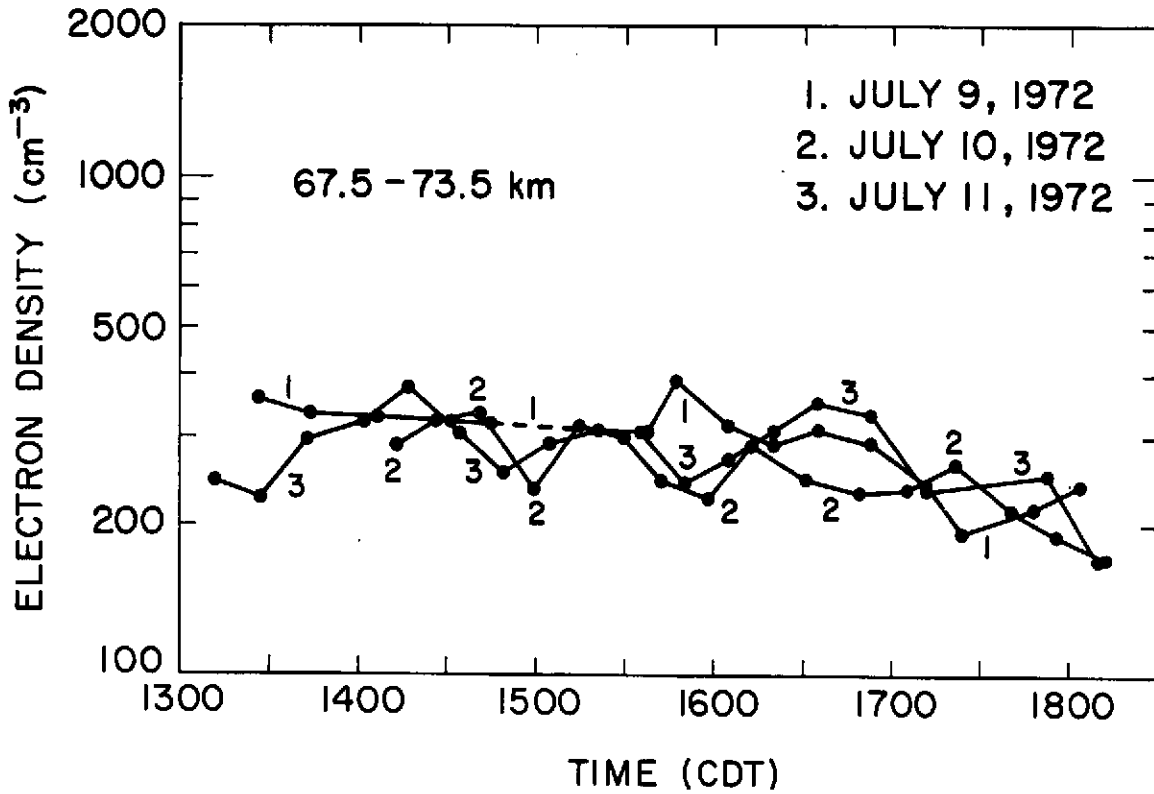


Figure 4.8 Median electron densities between 67.5 and 75 km.

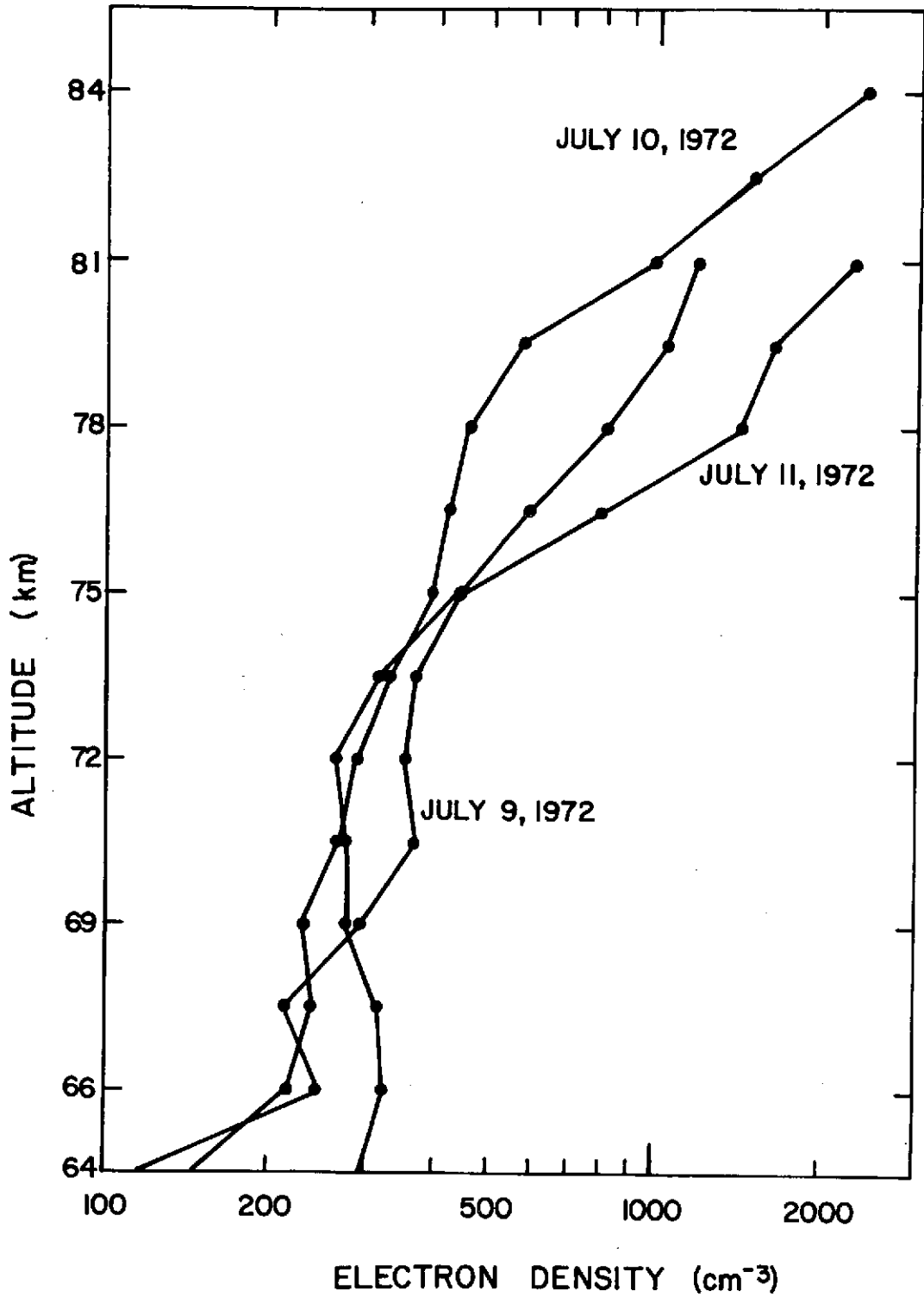


Figure 4.9 Median electron-density profiles between 1400 and 1500 CST.

to $[e]$ and α_{eff} . Equation (3.1)

$$q = \sigma_i(\text{NO}) [\text{NO}] I_\infty e^{-\tau} F_0 \quad (3.1)$$

where q = electron production rate in $\text{cm}^{-3} \text{sec}^{-1}$

$\sigma_i(\text{NO})$ = ionization cross-section of nitric oxide = $2 \times 10^{-18} \text{cm}^2$

$[\text{NO}]$ = number density of nitric oxide in cm^{-3}

I_∞ = incident Lyman-alpha flux at the top of the atmosphere = 3.1×10^{11}
photons $\text{cm}^{-2} \text{sec}^{-1}$

F_0 = the function of the unobscured solar disk

τ = optical depth

given by *Sechrist* [1966], was used to approximate the electron production rate and equation (3.2) was used to approximate the optical depth.

$$\tau = \sigma_a(\text{O}_2) [\text{O}_2] H \sec \chi \quad (3.2)$$

$[\text{O}_2]$ = number density of molecular oxygen in cm^{-3}

H = scale height

χ = solar zenith angle

Figure 4.10 shows the variation of q during the eclipse as compared to the variation without the eclipse. The electron production rates were used to obtain theoretical electron densities with α_{eff} being chosen to give the best fit to the experimental results. A value of 2×10^{-6} for α_{eff} was determined for the eclipse day between 75 and 82.5 km and 1.77×10^{-6} for the same height range on the control days. For the heights 78 to 87.5 km α_{eff} was found to be 8.46×10^{-7} . These values for α_{eff} are similar to ones given by *Mitra* [1968]. Figure 4.11 shows a comparison between the theoretical $[e]$ during the eclipse and without the eclipse using α_{eff} of 1.77×10^{-6} .

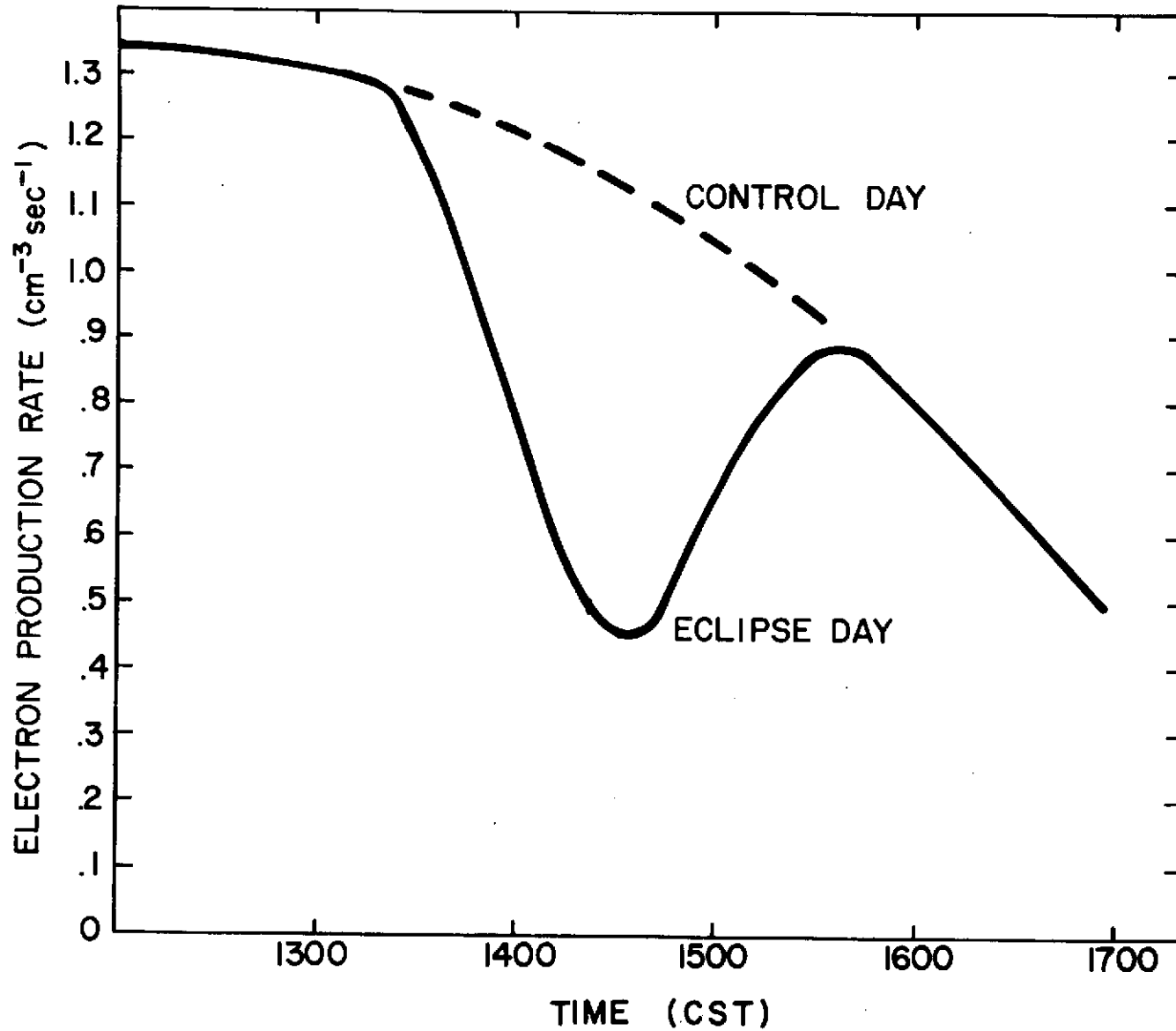


Figure 4.10 Electron production rate between 75 and 82.5 km during the eclipse and during the control days. The NO distribution used is from *Meira* [1971].

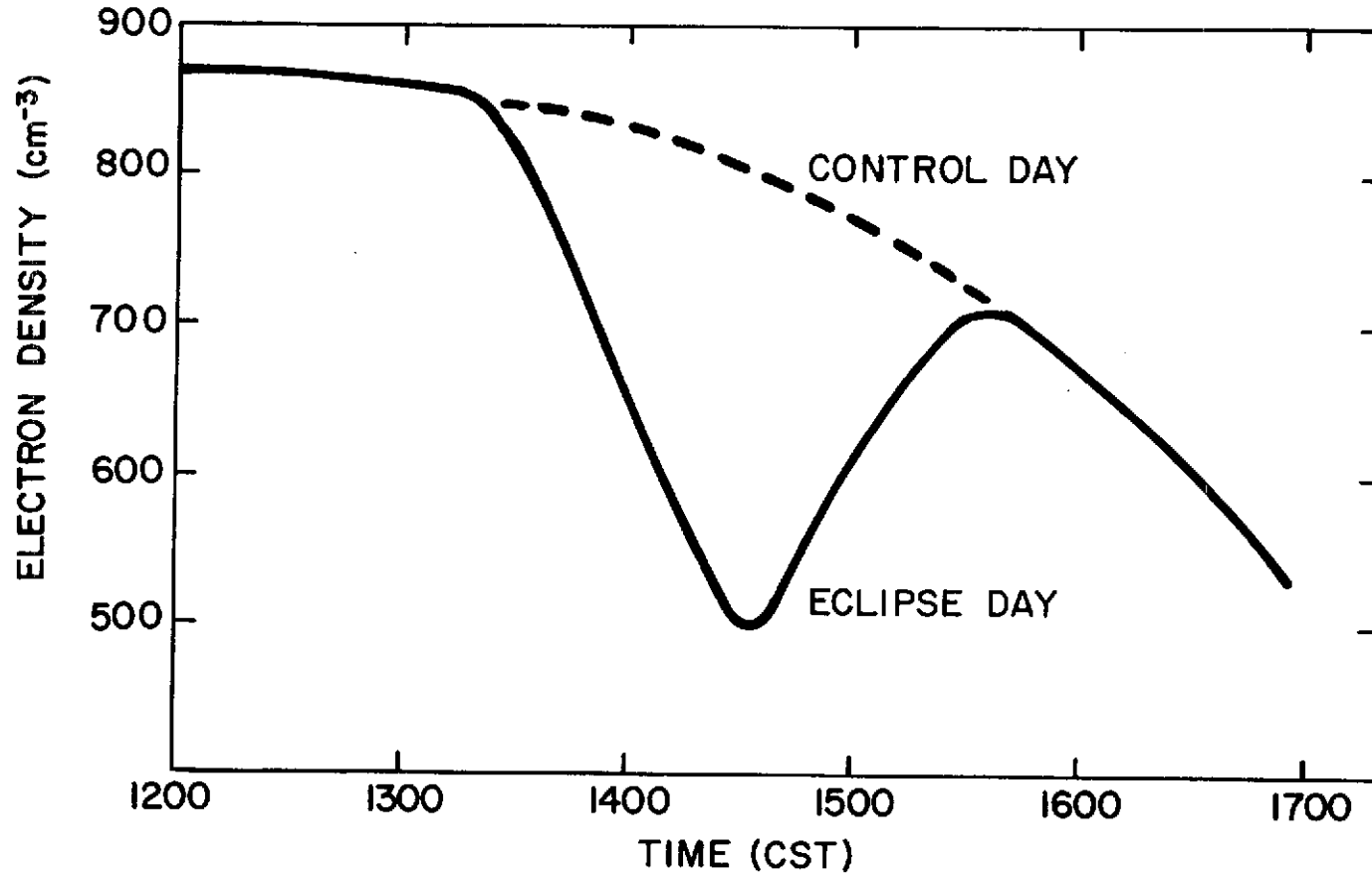


Figure 4.11 Theoretical electron densities between 75 and 82.5 km for eclipse and control day; calculated using an α_{eff} of 1.77×10^{-6} .

The electron density of the eclipse was divided by average electron density of the control data and compared to the obscuration function as seen in Figure 4.12. The comparison of the experimental $[e]$ during the eclipse and the theoretical $[e]$ without the eclipse using equation (2.4) was also made and is shown in Figure 4.13.

The electron density for July 9 shows a good correlation with the solar zenith angle (Figure 4.14) and was therefore divided by the theoretical $[e]$ to eliminate the effects of the solar zenith angle and to determine the variability of the experimental $[e]$ (Figure 4.15). The same comparison is made with the eclipse $[e]$ (Figure 4.15) and shows a similar but greater variability.

Generally, the eclipse electron densities show a decrease that is greater than expected from the equation (2.4). Other than the possibility that this is caused by variabilities due to inaccuracies in the experiment, there are three reasons why this may occur:

1. The obscuration function of the ionization source (Lyman- α) is different than the uniform-disk obscuration function used.
2. The α_{eff} increased during the eclipse. This could be caused by a change in the hydrated-ion composition between 75 and 81 km.
3. Loss by attachment is increased by the eclipse.

The electron-density profiles in Figure 4.10 show good comparison with the profile with 40% obscuration given in Figure 2.7 and with 60% obscuration shown in Figure 2.6. *Smith, et al.* [1965] described small changes below 70 km as the C-layer caused by cosmic rays which disappear as the eclipse reaches totality. The effect can be seen up to 69 km in Figure 4.9.

4.4 Summary

Comparing Figure 4.9 Figures 2.6 and 2.7, the electron-density profiles of this eclipse are similar to previous eclipses for the same obscuration. Generally, similar conclusions can be drawn. The difficulty in interpreting the

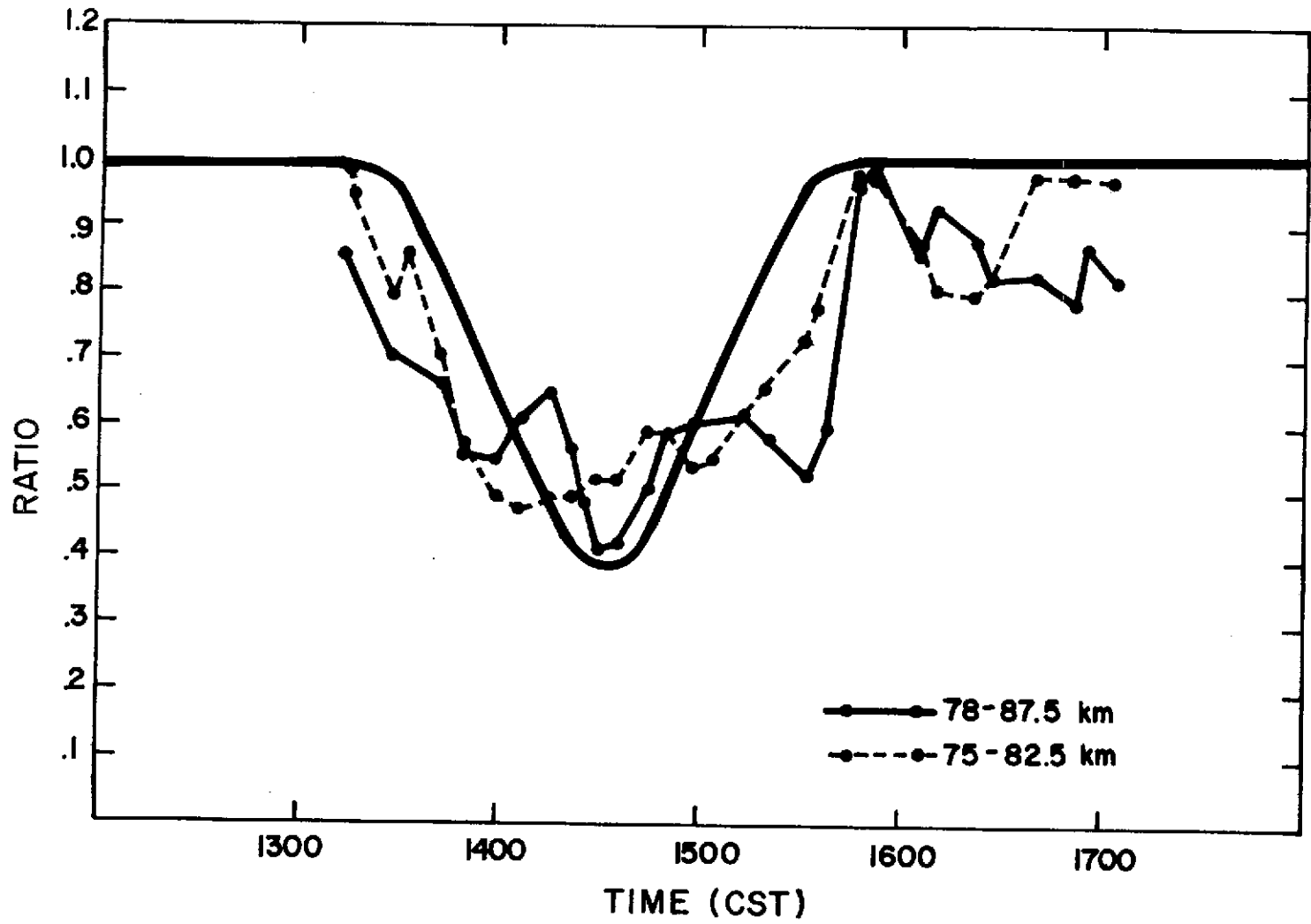


Figure 4.12 The ratio of electron densities for the average of the control day as compared to the unobscured sun.

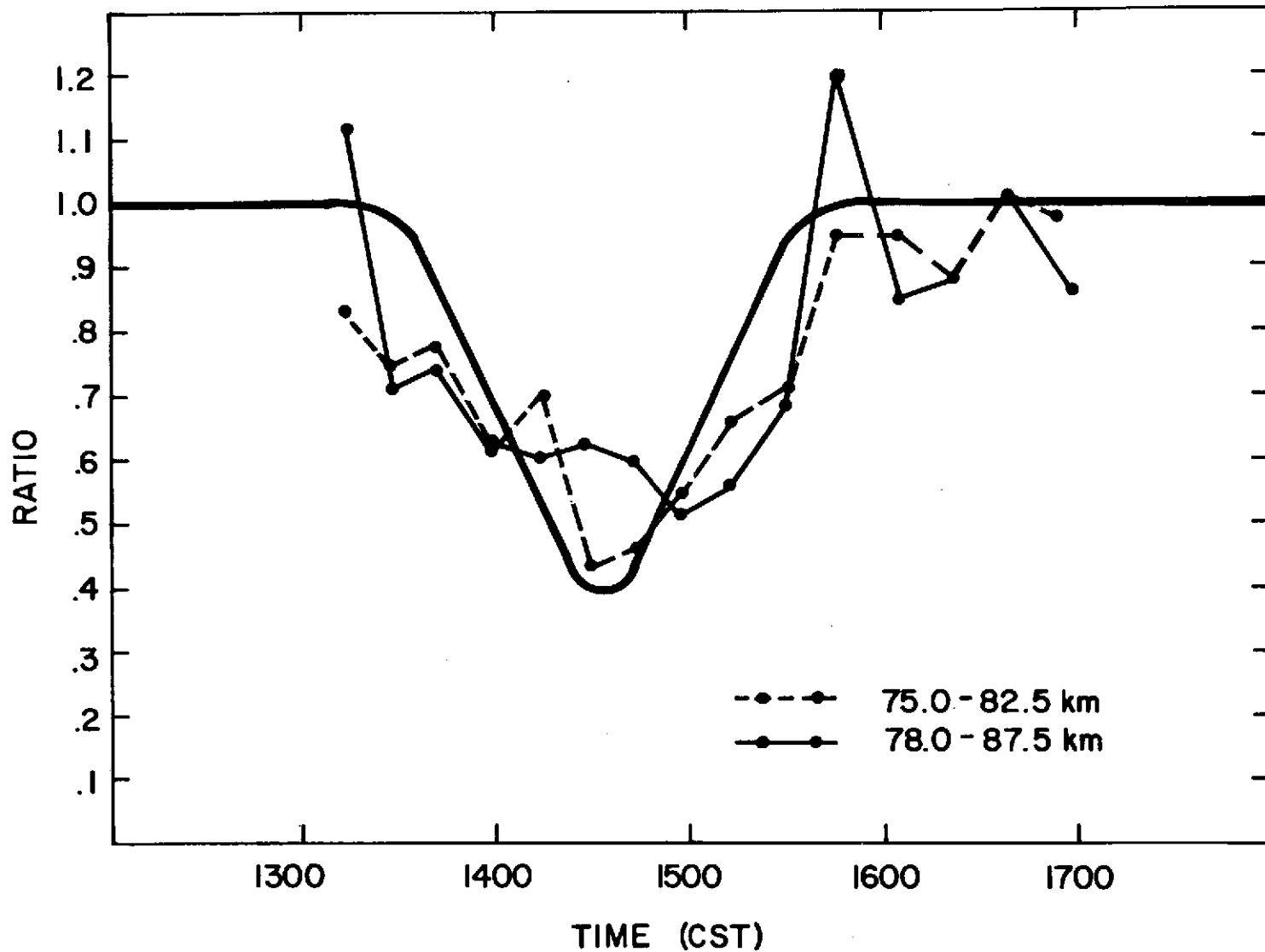


Figure 4.13 The graph of the ratios of the theoretical $[e]$ for the unobscured sun to the experimental $[e]$ for the eclipse as compared to the unobscured sun. The α_{eff} used for 75 to 82.5 km is 1.77×10^{-6} and for 78 to 87.5 km is 8.46×10^{-7} .

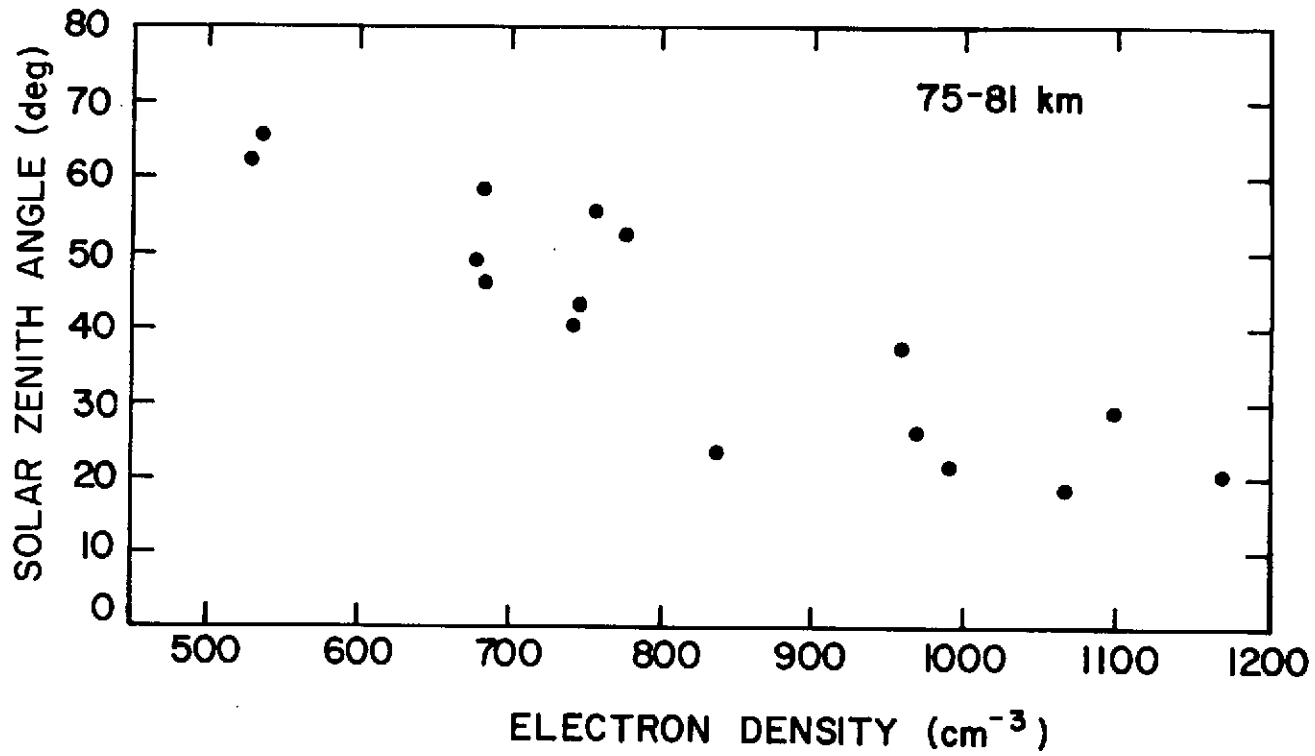


Figure 4.14 Scatter plot correlating the electron density for July 9, 1972 between 75 and 82.5 km to the solar zenith angle.

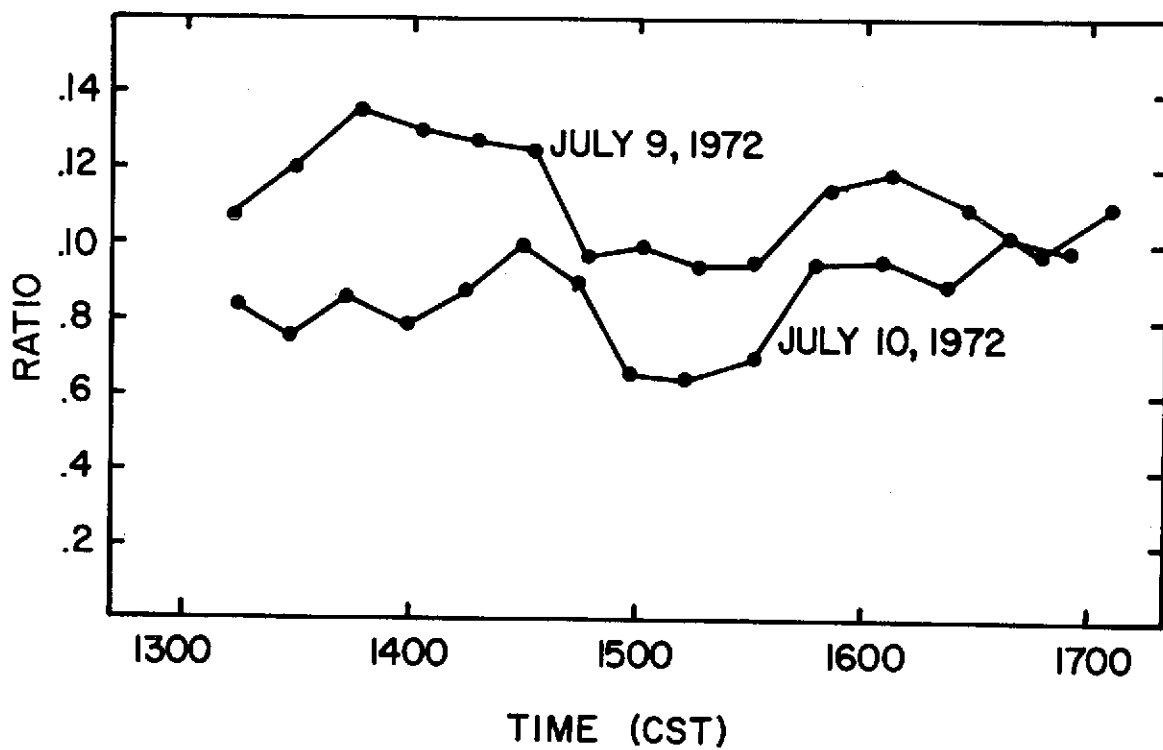


Figure 4.15 The graph of the ratio of theoretical electron densities to the experimental electron densities for July 9, 10, 1972.

results lies in the variation of the electron density of the eclipse with time. In Figures 4.12 and 4.13 a small decrease in electron density precedes the obscuration of the sun. An error of 20% can be expected due to the equipment and 20% error can be expected in the collision frequencies. Errors due to collision frequencies will cancel in Figure 4.12 but the errors due to the equipment will increase. For Figure 4.13, the reverse is true, but there are also errors due to the approximations made in equations (2.4), (3.1), and (3.2). With these possibilities of errors and observing that the ratio after the eclipse can get as low as .8 in Figure 4.12, the initial decrease can be interpreted as experimental error. The errors in Figure 4.13 can be seen in the variations in Figure 4.15.

No correlation could be seen between the X-ray flux and the electron density on the third day except during the X-ray burst period. Therefore, Lyman- α is assumed to be the main ionization source and the theoretical calculations were made on that assumption.

Of the three reasons for the large decrease in $[e]$, the effects due to changes in hydrated ions is the most likely. During the day electron loss by attachment is insignificant above 75 km. Since the obscuration of the sun was only 60% which corresponded to a production rate similar to that of 65° solar zenith angle, the loss process would still be by recombination.

The larger concentrations of Lyman- α on the solar disk were in the southern hemisphere and were not obscured and the intensity of 1-8 Å X-ray flux was too small to have any large effect. Therefore, the obscuration function of the ionizing source would have the same obscuration or less. This leaves the only possibility for the larger decrease in free electron as being due to changes in the hydrated ions.

5. CONCLUSIONS

The solar eclipse provides a good opportunity to study several processes of the D region and to develop its theoretical model. Accurate interpretation of the eclipse data is required to determine exactly the D -region ion production and loss processes, the variation of α_{eff} , formation of hydrated ions, and negative ion chemistry. A brief theory of the D -region chemistry is presented in Chapter 2 and used to analyze the data in Chapter 4. The equipment used in the collecting and processing of the partially reflected waves, as well as the refinements made in the collection process are given in Chapter 3. The newer partial-reflection system, discussed in Chapter 3, has been in use for the daily collection of data. Results from this newer system are given by *Denny and Bowhill* [1973]. This chapter reviews the results of the partial-reflection data taken during the eclipse and suggests further developments of the partial-reflection system.

5.1 *Review of Results*

The effect of the eclipse below 75 km is below the experimental errors. These errors are due to the variability of receiver gain caused by temperature fluctuations, the 40 μsec pulse width of the transmitter, inaccuracies in the collision frequencies, and inaccuracies in noise reduction. In comparing the $[e]$ profiles for July 9 and 10, 1972 in Figure 4.8, the beginning of the formation of a C layer can be seen resulting from cosmic rays. From Section 2.1, the main ionization source between 70 and 80 km is Lyman- α since the X-ray source effects were not observed below 81 km except when the X-ray flux increased above $1 \times 10^{-3} \text{ erg cm}^{-2} \text{ sec}^{-1}$.

The decrease between the electron density from July 9 and from July 10 is dependent on the height and is very marked between 79 and 81 km. Near 80 km,

this change in electron density is as much as 55% between the results of July 9 and 10, which was not expected according to equation (2.4). The most probably answer given in Chapter 4 is that it is due to an initial large decrease in hydrated positive ions which are the major ions between 75 and 80 km during the daytime (as seen in Figure 5.1 by *Krankowsky, et al.* [1972]).

The theoretical $[e]$ were used to compare with the experimental $[e]$ in Figure 4.13 to remove any electron density variability not due to the eclipse. The results in Figures 4.12 and 4.13 show unexpected initial decreases in $[e]$ prior to the eclipse and larger decreases than would be expected during the eclipse, but allowing for 20% error in these results, these variations are within the error limits. In general, there is good agreement with the data from *Smith, et al.*, [1927] and *Deeks* [1966].

5.2 Suggestions for Further Work

The present partial-reflection system has proved invaluable in presenting variations in electron densities diurnally and from day-to-day as presented by *Denny and Bowhill* [1973]. The system has several limitations, though. Either the signal-to-noise-ratio should be increased or the rates of data collection increased. Both of these changes would require alterations in the transmitter. By doubling the peak power of the transmitter, meaningful partial reflections could be obtained at lower altitudes without excessively disturbing the ionosphere due to the slow pulse rate as is done in the cross modulation experiment. By increasing the pulse rate, more data could be collected in the same interval of time, allowing for a more accurate statistical evaluation of the noise.

A new receiver has been built as mentioned in Chapter 3. The initial results obtained using it show an improvement in the results, but the problem of eliminating atmospheric noise remains. The main problem lies in defining the noise.

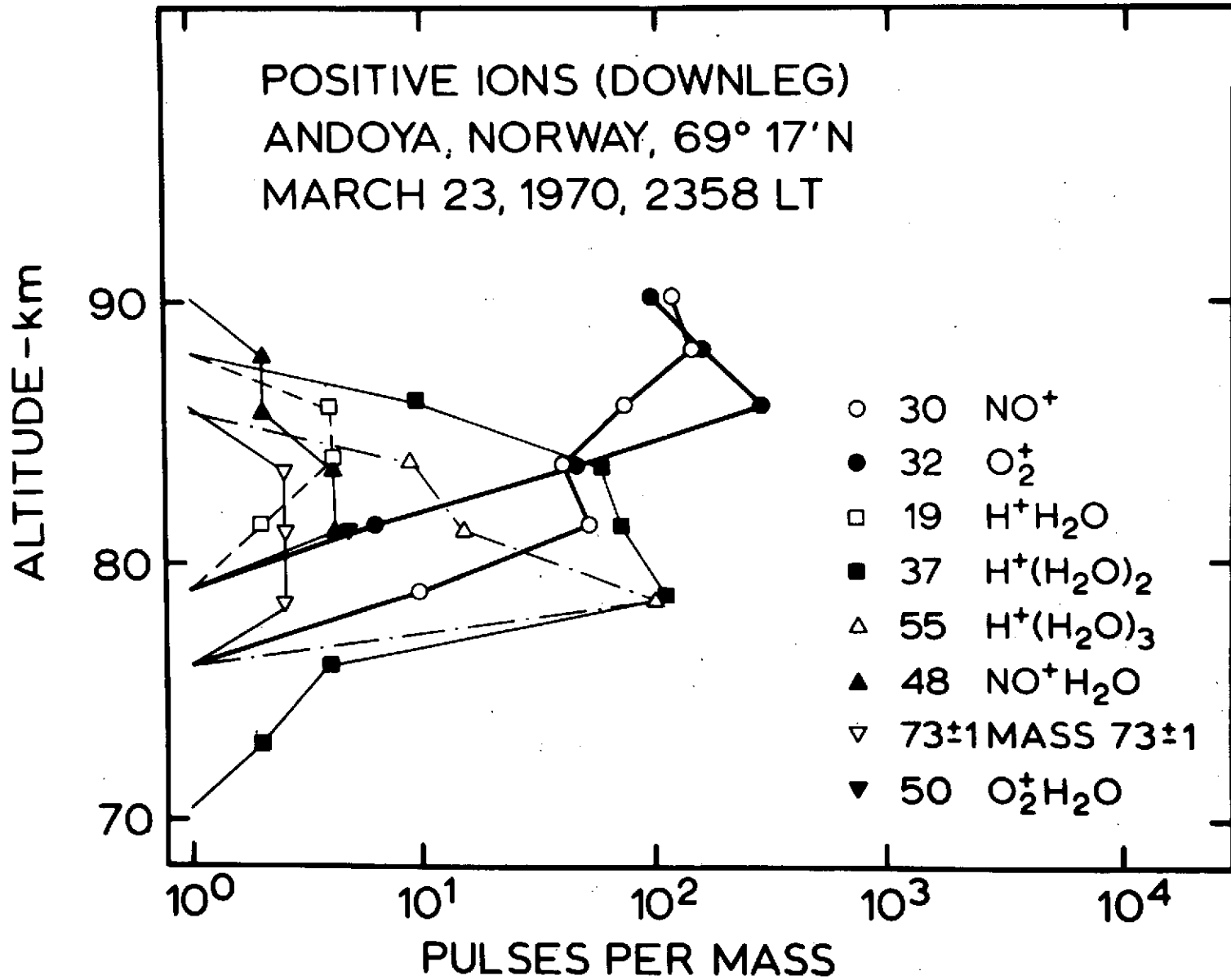


Figure 5.1 Rocket measurements of the positive-ion chemistry by Krankowsky, et al. [1972].

A study should be done on the specific types of noise received and the algorithms required to reject each. This would include receiving and storing noise on DECTape for later evaluation of the amplitude and phase.

A digital input/output would increase the efficiency of the collection and process. Presently the system requires the assistance of the operator every 3-1/2 minutes and uses one page of computer paper for every page of data. With a digital input/output, the computer could set the attenuators and control other switching which would free the operator for other tasks. This would also improve the usefulness of taking differential phase measurements as described by *Wiersma and Sechrist* [1972].

Using a line printer for outputting the data would allow for more sophisticated and complicated processing of data. This would also be required if the rate of collection is increased. To collect one file of data takes 3.5 minutes, to process one, about 45 sec, but to print out the results on the teletype and paper tape takes 2.6 minutes. Therefore, the processing would not be able to keep up with a faster collection unless the speed of printing the results increased.

REFERENCES

- Aikin, A. C. (1972), The relationship of theory and experiment in the *D* region, *J. Atmos. Terr. Phys.* 34, 1591-1599.
- Barth, C. A. (1966), Rocket measurement of nitric oxide in the upper atmosphere, *Planet. Space Sci.* 14, 623-630.
- Belrose, J. S. and M. J. Burke (1964), Study of the lower ionosphere using partial reflection, *J. Geophys. Res.* 69, 2799-2818.
- Birley, M. H. and C. F. Sechrist, Jr. (1971), Partial-reflection data collection and processing using a small computer, *Aeron. Rep. No. 42*, Aeron. Lab., Dep. Elec. Eng., Univ. Ill., Urbana-Champaign.
- COSPAR International Reference Atmosphere* (1965), North Holland Publishing Company, Amsterdam.
- Deeks, D. G. (1966), *D*-region electron distributions in middle latitudes deduced from the reflexion of long radio waves, *Proc. Roy. Soc. A291*, 413-437.
- Denny, B. W. and S. A. Bowhill (1973), *D*-region electron densities for the winter of 1972-73, *Aeron. Rep. No. 56*, Aeron. Lab., Dep. Elec. Eng., Univ. Ill., Urbana-Champaign
- Donahue, T. M. (1972), Positive ion chemistry of the *D* and *E* regions, *Radio Sci.* 7, 73-80.
- Edwards, B. (1972), Research in Aeronomy: October 1, 1972 - March 31, 1973, *Prog. Rep. 73-1*, Aeron. Lab., Dep. Elec. Eng., Univ. Ill., Urbana-Champaign.
- Fehsenfeld, F. C., A. L. Schmeltekopf, and E. E. Ferguson (1965), Correction in the laboratory measurement of the rate constant for $N_2^+ + O_2 \rightarrow N_2 + O_2^+$ at 300° K *Planet. Space Sci.* 13, 919-920.

- Fehsenfeld, F. C. and E. E. Ferguson (1969), Origin of water cluster ions in the D-region, *J. Geophys. Res.* 74, 2217-2225.
- Fejer, J. A. and R. W. Vice (1959), An investigation of the ionospheric D region, *J. Atmos. Terr. Phys.* 16, 291-306.
- Ferguson, E. E. (1971), D-region ion chemistry, *Rev. Geophys. Space Phys.* 9, 997-1008.
- Gardner, F. F. and J. L. Pawsey (1953), Study of the ionospheric D-region using partial reflections, *J. Atmos. Terr. Phys.* 3, 321-344.
- Good, A., D. A. Durden, and P. Kebarle (1970), Ion-molecule reactions in pure nitrogen and nitrogen containing traces of water at total pressures .5-4 torr. kinetics of clustering reactions forming $H^+(H_2O)_n$, *J. Chem. Phys.* 52, 222-229.
- Gregory, J. B. (1956), Ionospheric reflections from heights below the E region *Australian J. Phys.* 9, 324-342.
- Henry, G. W., Jr. (1966), Instrumentation and preliminary results from shipboard measurements of vertical incidence ionospheric absorption, *Aeron. Rep. No. 13*, Aeron. Lab., Dep. Elec. Eng. Univ. Ill., Urbana-Champaign.
- Huffman, R. E., D. E. Paulsen, J. C. Larrabee, and R. B. Cairns (1971), Decrease in D-region $O_2(^1\Delta_g)$ photoionization rates resulting from CO_2 absorption, *J. Geophys. Res.* 76, 1028-1038.
- Hunt, B. G. (1965), A theoretical study of the changes occurring in the ozonosphere during a total eclipse of the sun, *Tellus XVII*, 516-523.
- Hunten, D. M. and M. B. McElroy (1968), Metastable $O_2(^1\Delta)$ as a major source of ions in the D region, *J. Geophys. Res.* 73, 2421-2428.
- Keneshea T. J. (1967), A technique for solving the general reaction-rate equations in the atmosphere, *Environmental Res. Papers No. 263*, Air Force Cambridge Research Laboratories, Bedford, Massachusetts.

- Krankowsky, D., F. Arnold, H. Wieder, J. Kissel, and J. Zähringer, (1972), Positive-ion composition in the lower ionosphere, *Radio Sci.* 7, 93-98.
- Lauter, E. A., and R. Knuth (1967), Precipitation of high energy particles into the upper atmosphere at medium latitudes after magnetic storms, *J. Atmos. Terr. Phys.* 29, 411-417.
- Lodato and Mechtly (1971), Rocket measurements of electron collision frequency, *Aeron. Rep. No. 45*, Aeron. Lab., Dep. Elec. Eng., Univ. Ill., Urbana-Champaign.
- Mechtly, E. A., S. A. Bowhill, L. G. Smith, and H. W. Knoebel, (1967), Lower ionosphere electron concentrations and collision frequency from rocket measurements of Faraday rotation, differential absorption, and probe current, *J. Geophys. Res.* 72, 5239-5245.
- Mechtly, E. A., C. F. Sechrist, Jr., and L. G. Smith (1972), Electron loss coefficient for the D-region of the ionosphere from rocket measurements during the eclipses of March 1970 and November 1966, *J. Atmos. Terr. Phys.* 34, 641-646.
- Meira, L. G., Jr. (1971), Rocket measurements of the upper atmospheric nitric oxide and their consequences to the lower ionosphere, *J. Geophys. Res.* 76, 202-212.
- Mitra, A. P. (1966), An ionospheric estimate of nitric oxide concentration in the D-region, *J. Atmos. Terr. Phys.* 28, 945-955.
- Mitra, A. P. (1968), D-region processes in non-polar latitude, *J. Atmos. Terr. Phys.* 30, 1065-1114.
- Mitra, A. P. and J. N. Rowe (1972), Ionospheric effects of solar flares - VI changes in D-region ion chemistry during solar flares, *J. Atmos. Terr. Phys.* 34, 795-806.

- Narcisi, R. S. and A. D. Bailey (1965), Mass spectrometric measurements of the positive ions at altitudes from 64 to 112 kilometers, *J. Geophys. Res.* 70, 3687-3700.
- Pearce, J. B. (1969), Rocket measurements of nitric oxide between 60 and 96 kilometers, *J. Geophys. Res.* 74, 853-861.
- Pirnat, C. R. and S. A. Bowhill (1968), Electron densities in the lower ionosphere deduced from partial reflection measurements, *Aeron. Rep. No. 29*, Aeron. Lab., Dep. Elec. Eng., Univ. Ill., Urbana-Champaign.
- Radicella, S. M. and D. W. Stowe (1970), D-region ion chemistry, *Aeron. Rep. No. 38*, Aeron. Lab., Dep. Elec. Eng. Univ. Ill., Urbana-Champaign.
- Reynolds, D. A. and C. F. Sechrist, Jr. (1970), Measurement of average electron density between 75 and 80 kilometers, *Aeron. Rep. No. 36*, Aeron. Lab., Dept. Elec. Eng., Univ. Ill., Urbana-Champaign.
- Sears, R. D. (1972), Analysis of the 1966 solar eclipse data, *DNA2863F, LMSC D246526*, Lockheed Palo Alto Res. Lab., Calif.
- Sechrist, C. F., Jr. (1966), A theory of the winter absorption anomaly at middle latitudes, *J. Atmos. Terr. Phys.* 29, 113-136.
- Sechrist, C. F., Jr. (1972), Theoretical models of the D-region, *J. Atmos. Terr. Phys.* 34, 1565-1589.
- Sen, H. K. and A. A. Wyller (1960), On the generalization of the Appleton-Hartree magnetoionic formulas, *J. Geophys. Res.* 65, 3931-3950.
- Shimazaki, T. and A. R. Laird (1970), A model calculation of the diurnal variation in minor neutral constituents in the mesosphere and lower thermosphere including transport effects, *J. Geophys. Res.* 75, 3221-3235.
- Shimazaki, T. and A. R. Laird (1972), Seasonal effects on distributions of minor neutral constituents in the mesosphere and lower thermosphere, *Radio Sci.* 7, 23-44.

- Smith, L. G., C. A. Accardo, L. H. Weeks, and P. J. McKinnon (1965), Measurements in the ionosphere during the solar eclipse of 20 July 1963, *J. Atmos. Terr. Phys.* 27, 803-829.
- Somayajulu, Y. V. and M. B. Avadbanula (1972), Rocket measurements of D-region electron densities at the equator, *Indian J. Radio Space Sci.* 1, 81-83.
- Thomas, L. (1971), The lower ionosphere, *J. Atmos. Terr. Phys.* 33, 157-195.
- Turco, R. P. and C. F. Sechrist, Jr. (1970), An investigation of the ionospheric D region at sunrise, *Aeron. Rep. No. 41*, Aeron. Lab., Dep. Elec. Eng. Univ. Ill., Urbana-Champaign.
- U. S. Standard Atmosphere Supplements* (1966), U. S. Government Printing Office, Washington, D. C.
- U. S. Department of Commerce (1972), *Solar-Geophysical Data* 337, part I, 58-59.
- U. S. Department of Commerce (1973), *Solar-Geophysical Data* 341, part II, 48-49.
- Whitten, R. C. and I. G. Poppoff (1971), *Fundamentals of Aeronomy*, John Wiley and Sons, Inc., New York.
- Wiersma, D. J. and C. F. Sechrist, Jr. (1972), Differential phase measurements of D-region partial reflections, *Aeron. Rep. No. 47*, Aeron. Lab. Dep. Elec. Eng. Univ. Ill., Urbana-Champaign.

APPENDIX

```

.TITLE DLOGFI
/ DLOGFI IS A COMBINATION OF ALL THE PARTIAL REFLECTION
/ MACRO PROGRAMS USED IN COLLECTING AND PROCESSING DATA. THE
/ PROGRAMS CONTAINED IN THIS VERSION ARE:
/INITIALIZATION: INTIM STARTS THE TIME
/ DLOGF INITIALIZES COLL.& PROC. PARAM.
/ READM READS UNFORMATTED CHAR. FROM TTY
/ TOD INCREMENTS THE TIME OF DAY
/CALIBRATION: RADC READS SAMPLES FROM A/D CONVERTER
/ FOR FORTRAN PROGRAMS
/ TTM WRITES LIN. TABLE OUT ON DISK
/COLLECTION: BEGIN MAIN COLL. PROG.--A REAL TIME
/ SUBROUTINE API LEVEL 6
/ WT1 GIVES TIME TO LOWER API LEVELS
/ RSUB SETS UP DATA PACK & CHECKS NOISE
/ --R.-T. SUB. AT API LEVEL 5
/ PAC PACKS DATA DOUBLE
/ DTRANS WRITES DATA ON STORAGE DEVICE
/ CKCNT CHECKS FOR ENOUGH DATA
/ CHECK SEES IF AX & AD IS IN THE RIGHT
/ ORDER AND ARE COLL. IN PAIRS--API 5
/ LIN LIN. DATA,NEG #=0,& CHECKS NOISE
/ ADREAD PREPARES A/D CONVERTER READ
/ ADINT A/D INTERRUPT SERVICE ROUTINE
/PROCESSING: CHNG INIT. DEV. & CHECKS FOR COLL.FIE
/ CONTR WAITS FOR FILE TO BE COLLECTED
/ JAITI ALLOWS TIME FOR BACKGROUND
/ CKCOL CHECKS FOR UNWANTED COLL.STOP
/ SWIT USED TO SWITCH DISKS
/ DUMPT READS DATA, UNPACKS IT,& PUTS
/ IT INTO A FORTRAN ARRAY
/
/ DATA IS COLLECTED ALTERNATELY ORDINARY AND EXTRAORDINARY
/AS DESCRIBED BY HIRLEY (AERONOMY REPORT 42). THE TIMING OF THE
/ DATA IS DETERMINED BY AN EXTERNAL ENCODE PULSE. THE MODE OF THE
/ DATA INPUTED INTO THE COMPUTER IS DETERMINED BY A TIMING PROGRAM
/ (CHECK) SET UP BY D. WARD. ESSENTIALLY HOW IT WORKS IS AFTER A
/ FRAME OF DATA HAS BEEN READ IN THE COMPUTERS CLOCK IS SET FOR 9/64
/ OF A SECOND (9 PULSES). IF NO OTHER DATA IS READ IN BEFORE THE
/ TIME EXPIRES THE DATA FRAME WAS EXTRAORDINARY MODE AND IS REJECTED.
/ OTHERWISE BOTH FRAMES ARE ACCEPTED. THIS CHECK IS MADE AT THE
/ BEGINNING, AFTER EACH DATA TRANSFER, AND WHENEVER THE COLLECTION
/ IS RESTARTED OR AN ERROR CONDITION EXIST. THE PROGRAM IS SET SO
/ AS TO NOT OVER MAXIMUM STORAGE ON THE DISK.
/
/*****
/
TTI=4 /TELETYPE IN
TTO=6 /TELETYPE OUT
TBI=10 /.DAT SLOT OF LIN. TABLE
OUTPT=2 /PLACE TO STORE DATA
OUTPT2=1 /SECOND PLACE TO STORE DATA
DATIN=5 /.DAT SLOT TO READ DATA
DATIN2=3 /SECOND .DAT SLOT TO READ DATA
OUT=1 /OUTPUT TO I/O DEVICE
IN=0 /INPUT FROM I/O DEVICE
ASC=2 /TYPE OF I/O MODE
IA=3 /TYPE OF I/O MODE
DUMP=4 /TYPE OF I/O MODE
SKAR=5 /# OF DATA TO BE DELETED
TNSAM=37 /# OF SAMPLES TO BE READ IN
NSAM=TNSAM-SKAR /# OF SAMPLES PER FRAME TO BE STORED
NSAMP=NSAM/2+1 /SIZE OF ONE FRAME PACKED DOUBLE
DATBLK=TNSAM+2 /SIZE OF INITIAL DATA BLOCK
DTBLK=374 /SIZE OF 1 BLOCK OF STORAGE
DATSTR=DTBLK/NSAMP /# OF FRAMES FOR 1 BLOCK OF STORAGE
RBLK=10 /SIZE OF BLOCK FOR TTI READ
MNCC=-11 /COUNT. FOR # OF NOISE FOR MAX. NOISE
NFFC=-6 /MINUS (#+1) OF NOISE PER FRAME
NDPC=NSAM/2 /# OF DATA PER FRAME
DPR=NDPC+1 /RESET FOR POINTER TO DATA
MXFPD=350000/NSAMP /MAX. # FRAMES PER DEVICE
NOFDK=1 /# OF DISKS TO BE USED
/
/ GLOBL CHNG,.DA,.AD,DUMPT,DLOGF,CKCOL,CONTR,PROC,RADC,TTM
/ GLOBL TBFORL,PP7,VFPP
/ IODEV 1,2,3,4,5,6,10 /.DAT TO BE USED
/*****
/ THE FOLLOWING SUBROUTINE IS USED TO PREPARE THE
/ COLLECTION AND PROCESSING PROGRAMS FOR MANIPULATION OF DATA.
/ THE VARIABLES OF THE MACRO PROGRAMS ARE STORED IN THIS SUBROUTINE.
/ THEREFORE AFTER THIS SUBROUTINE IS EXECUTED IT IS WRITTEN OVER
/ AND SHOULD NOT BE REENTERED (FOR IT WILL NOT EXIST).

```

```

/
DLOGF 0
TCR JMS* .DA /GET SUBROUTINE PARAMETERS' ADDRESSES
SUM4 JMP .+4 /JUMP AROUND PARAMETER LIST
DB 0 /ADDR. OF THE ADDR. OF THE DB SETTINGS
DBC 0
TC3 0
RMSG1 LAC MSG1+3 /SET UP RESET FOR TELETYPE
TC1 DAC RMSG1 / MESSAGE
DSTOR JMP ONG4 /GO TO NEXT EXECUTABLE STATEMENT
CTT1 LAW NPFC /COUNTS # OF NOISE PER FRAME
CTT2 LAW MNCC /COUNTS GROUPS OF NOISE FOR MAX. NOI.
CTT3 LAW -2 /USED TO SWITCH STORAGE DEVICES (IF
CTT4 LAW -2 / NEEDED) FOR COLL. AND PROC.
CTT5 LAW NPFC /USED TO SKIP AROUND "SKAR" DATA #'S
CTT6 LAW -2 /USED TO SWITCH BUFFERS IN COLL.
CTT7 0 /TELLS # OF UNWANTED STOPS IN COLL.
CTT8 LAW -DATSTH /COUNTS # OF DATA #'S PER BUFFER
CTT9 LAW -1 /ALLOWS PROC. TO READ DATA
CTT10 0 /TELLS PROC TO RESTART A FILE 1
CTT11 0 /TELLS PROC. END OF RUN
ALF 51004 /USED TO PUT A LINE FEED IN THE
20100 / MIDDLE OF A FORTRAN OUTPUT LINE
TST1 0 /SAVE ID TO TEST FOR COLL. STOPAGE
SAVAC 0 /SAVES AC DURING PROC. OPERATION
SAVEAC 0 /SAVES AC FOR API LEVEL 6
TC10 LAW -2 /USED TO UNPACK 2 WORDS
CN1 0 /STORES # OF FRAMES FOR ONE FILE
CN2 0 /STORES # OF FRAMES FOR ONE DISK
CNT 0 /ODD FOR 0-FRAMES, EVEN FOR X-FRAMES
CNT2 0 /SET TO -1 FOR 0-FRAMES
DUM1 0 /GETS & STORES 1 LINEARIZED DATA WORD
DUM2 0 /STORES THE MAX. NOISE IN A FRAME
DUM3 0 /STORES THE SUM OF THE NOISE PER FRAME
DUM4 0 /STORES THE MAX. OF 45 NOISE SAMPLES
DUM5 0 /STORES THE SUM OF 45 NOISE SAMPLES
MAX4 777 /STORES THE SMALLEST MAX. OF "DUM4"
TBUF .DSA BUF1 /STORES NAME OF BUF. TO BE OUTPUTED
IDCOU 0 /STORES THE ID NUMBER
TRANF 0 /DATA IS TRANSFERED WHEN NON-ZERO
COUNTP LAW -NDPC /COUNTS # OF DATA PER FRAME TO PACK
BPOINT .DSA BUF /POINTER FOR "BUF"
COUNT LAW -NDPC /COUNTS # OF POINTS TO UNPACK
ER2 0 /STORES TIMING ERROR FOR A-D CONVERTER
INFLAG 0 /TIMING ERROR FLAG
INSUB 0 /STORES R.-T. SUB. & API LEVEL FOR A/D
TIME 0 /STORES TIME (RESET EVERY .1 MIN.)
TIMR 0 /STORES TIME USED IN DATA HEADING
SVAC 0 /SAVES AC DURING TIMER ROUTINE (TOD)
MIN 0 /STORES THE MIN. IN DURING EACH HOUR
MCNTR 0 /STORES # OF MIN. OF RUN
LETCRG .DSA NBLK /ADDRESS OF SUB. "READM" BUFFER
BUF11 .DSA BUF1 /ADDRESS OF BUF1
BUF21 .DSA BUF2 /ADDRESS OF BUF2
BUF31 .DSA BUF2 /ADDRESS OF BUF3
SEWN 37000 /133151 DECIMAL
.DEC
MMIN 599 /MAX # OF .1 MIN.'S IN 1 HOUR
DRI 23599 /MAX. AMOUNT OF TIME IN ONE DAY
DHR 1000-599 /USED TO INCREMENT TIME BY 1 HOUR
.OCT
DBP .DSA DB0 /POINTER FOR THE DB'S
DB0 0 /DB'S STORAGE FLOCK
30540 /ASCII DEFAULT SETTINGS
31152 / ARE 0,10,25 (DEC)
LAW -1 /TELLS END OF DB SETTINGS
0
ST1 .BLOCK 1000 /LIN. TABLE
POINT .DSA SEWN /POINTER FOR COLLECTION BUFFERS
BUF1 .BLOCK DTBLK-323 /FIRST COLLECTION BUFFER
STT .SIXBT "TABLEADAT" /NAME OF FILE TO BE READ IN
MSG7 2000
0
LAST .ASCII "SET CONSOLE SWITCHES, TURN ON PULSER AND ENCODE PU"
.ASCII "LSE"<15>
MSG8 2000
0
.MSG11 "DISCONNECT WIRE FROM PULSED OSCILLATOR"<15>
MSG9 2000
0
.MSG12 "SET ATTENUATOR TO 00DB"<15>
MSG10 2000
0
.MSG13 "TYPE # OF HOURS FOR THE RUN AND C.R."<15>
MSGDB 2000
0

```



```

.ERRCAL .ASCII "DB SETTING"<15>
2000
0
.ASCII "ERROR IN CALIBRATION TABLE"<15>
DMV1 JMP BLK2 /USED TO DELETE COLL. READ
DMV2 204 /LOC. TO DETERMINE WHICH TTY CONFIRMS IC
DMV3 177 /LOC. USED TO ALLOW SHARE
DMV4 116 /ADDR. OF LOC. OF FOREGND .DAT SLOT 0
DMV5 117 /ADDR. OF LOC. OF BACKGND .DAT SLOT 0
DMV6 113 /ADDR. OF THE .IOIN TABLE
SRP0 0 /USED TO SURPRESS CALIBRATION
FDAT1 0 /THESE LOCATIONS ARE USED
FDAT2 0 / TO STORE THE
FDAT3 0 / FOREGROUND .DAT SLOT
FDAT5 0 / ADDRESSES FOR 1,2,3,4,5
FDAT10 0 / AND 10
MINEQC .DSA MINEQ /USED TO DETERMINE LENGTH OF RUN
.DEC
MINEQ 6000 /CONVERTS INPUT NUMBERS TO
600 / THE EQUIVALENT
60 / BCD # OF MINUTES
.OCT
CNTR1 0 /USED TO DETERMINE MULTIPLES OF 10
DBP1 0 /TEMP. STORAGE FOR DB SETTING
DBP2 0 /TEMP. STORAGE FOR DB IN ASCII CODE
DBCNT1 0 /COUNTER FOR # OF DB TO BE USED
DBCNT2 LAW -2 /ALLOWS NO MORE THAN 2 DIGIT DB'S
DBCNT3 LAW -4 /ALLOWS ONLY 4 DB SETTINGS
CODE1 777723 /-55 TO CECK FOR CARRIAGE RETURN
CODE2 .DSA MINEQ*2 /FOR LESS THAN 10 HOUR RUN
CODE3 10060 /ASCII SPACE AND A ZERO
CODE4 15 /CHECK FOR CARRIAGE RETURN
CODE5 LAW -2 /CHECKS FOR MULTI-DB SETTINGS
CODE6 20140 /ASCII DEFAULT DB FOR 0 DB SETTING
CODE7 7700 /USED TO FIND THE DEVICE # IN .IOIN
CODE8 777766 /-12 TO FIND BCD # MULTIPLE OF 10
CODE9 72 /SETS THE DB SETTING
CODE10 60 / TO ASCII CODE
CODE11 11 /INITIALIZES .IOIN TABLE POINTER
CODE12 12 /USED TO DETERMINE VALUES OF ASCII #'S
CODE13 1200 /LOOKS FOR DEVICE # 5 (DISK)
CODE14 1000 /LOOKS FOR THE DEVICE # 4 (DECTAPE)
CODE15 160000 /USED TO FIND THE UNIT # FOR A DEVICE
CODE16 16 /CHECKS FOR DECIMAL POINT
CODE17 LAW -32 /USED TO CHECK FOR NONNUMBER ASCII CHAR.
CODE18 -MXFPD /SETS MAXIMUM AMOUNT COLL. ON A DISK
CODE19 .DSA DECN /INIT. MULTIPLIERS FOR BCD #'S
CODE20 .DSA MINEQ /INIT. MULT. TO CONVERT TIME TO MIN
CODE21 NOFDK /SETS UP # OF DISKS TO BE USED
CODE22 20000 /CHECKS FOR UNIT 1
CODE23 60000 /CHECKS FOR UNIT 3
CODE24 -310000 /CHECK FOR ASCII "1" IN DATE
CODE25 -14000 /USED TO CHECK FOR ASCII "5" TO "7"
CODE26 3760 /USED TO GET SECOND CHAR. IN ASCII WORD
CODE27 1400 /CHECKS FOR AN ASCII ZERO
CODE28 360000 /USED TO CHECK FOR AN ASCII "A"
CODE29 -10000 /USED TO INCREMENT 2 ASCII LETTERS
CODE30 -20000 /USED TO INCREMENT 4 ASCII LETTERS
CODE31 -4000 /USED TO INCREMENT 1 ASCII LETTER
CODE32 20200 /CHECKS FOR AN ASCII "A"
CODE33 7 /USED TO MASK ALL HIGHER BITS
CODE34 3 /CHECKS FOR FIRST HALF ASCII "Y"
CODE35 -27340 /DEFAULT VALUE FOR TIME OF DAY (20 HR.)
CODE36 -100000 /CHECKS FOR UNITS LESS THAN 4
CODE37 101204 /DUMMY CODE FOR .IOIN TABLE
CODE38 402574 /SECOND WORD FOR .IOIN TABLE
CODE39 1600 /LOOKS FOR DEVICE #7 (P. PUNCH)
CODE40 200300 /USED TO GET CONSOLE SWITCH 1
/.....
.EJECT
/#####
/ BEIENNING OF THE PARTIAL REFLECTION PROGRAMS.
/ THIS PART INITIALIZES THE .DAT SLOTS, SETS UP NEW
/DIRECTORIES, ALLOWS SHARE MODE, AND SETS AND INITIATES THE
/CLOCK TO GIVE THE TIME OF DAY.
/
INTIM .INIT TTO,OUT,INTIM /TTY OUT
.INIT TTI,IN,INTIM /TTYIN
LAC+ DMV4 /GET ADDR. OF FOREGND .DAT SLOTS
IAC
DAC FDAT1 / ONE

```

```

IAC
DAC   FDAT2       /   TWO
IAC
DAC   FDAT3       /   THREE
IAC
IAC
DAC   FDAT5       /   FIVE
TAD   CODE34
DAC   FDAT10      /   EIGHT
LAC+  FDAT2       /SET .DAT SLOT 5 EQUAL
DAC+  FDAT5       /   TO .DAT SLOT 2
LAC+  FDAT3       /SET .DAT SLOT 1 EQUAL
DAC+  FDAT1       /   TO .DAT SLOT 3
.INIT  OUTPT,OUT,INTIM /PREPARE FIRST STORAGE DEVICE
.CLEAR OUTPT
LAW   -2          /ARE TWO STORAGE
TAD   CODE21      /   DEVICES
SPA   /          /   REQUIRED ?
JMP   NDK11      /NO, OMIT FOLOWING CODE
LAC   MDUM1      /YES, CLEAR JUMP AROUND INSTRUCTIONS
DAC   NDK2       /   TO USE A SECOND DISK

DAC   NDK3       /   FOR STORAGE
.INIT  DATIN2,OUT,INTIM /PREPARE SECOND DEVICE FOR USE
.CLEAR DATIN2    /REMOVE ALL FILES FROM IT
JMP   .+3

NDK11 DZM+  FDAT1    /CLEAR .DAT SLOTS FOR THE
      DZM+  FDAT3    /   SECOND STORAGE DEVICE
      LAS   /        /GET THE VALUE OF CONSOLE
      AND   CODE40   /   SWITCH #1
      DAC   SRP0     /SET ADDR. TO THIS VALUE
      LAW   -1
      DAC+  DMV3     /ALLOW SHARE
      DZM+  DMV2     /SET IT0 TO ACKNOWLEDGE %C
      JMP   RDERR    /JUMP AROUND SECOND BUFFER

/.....
BUF2  .BLOCK DTBLK-350 /SECOND COLLECTION BUFFER
/.....
RDERR LAC   CODE19   /INITIALIZE BCD
      DAC   HRC      /   POINTER
      ISZ  CTM1      /HAS FOUR NUMBERS BEEN
      SKP  /          /   READ
      JMP  OTLP      /YES, EXIT FROM ROUTINE
      .WRITE TTO,ASC,MSGT,0 /NO, ASK FOR TIME
      .WAIT TTO
      JMS  READM     /READ IN TIME
DT     0           /CONTAINS THE ADDR. OF CHAR. READ IN
      LAW   -5       /INITIALIZE COUNTER TO EXIT THE
      DAC   CTM1     /   ROUTINE AFTER FIVE #'S
      LAW   -3       /INITIALIZE COUNTER TO GET
      DAC   CTM2     /   THE MINUTES
      DZM   MIN      /INITIALIZE LOG. THAT SAVE
      DZM   HR       /   THE TIME AND MINUTES
      JMS  CHKN      /GET NEXT NUMBER
      JMS+ .AD        /MULTIPLY BY POWERS OF TEN
      LAC+ HRC       /   TO GET BCD EQUIVALENT
      DAC   SAV2     /SAVES THE MINUTES
      TAD   HR       /SETS UP THE NUMBERS READ IN
      DAC   HR       /   AS THE PRESENT TIME
      ISZ  CTM2     /IS THE NUMBER PART OF THE MIN. ?
      JMP  JPAR      /NO, FIRST TWO #'S ARE THE HOURS
      LAW   -1       /RESET COUNTER TO GET ALL THE
      DAC   CTM2     /   MIN. (REST OF THE #'S)
      LAC   SAV2     /GET MINUTES AND
      TAD   MIN      /   SET INTO AN ADDR.
      DAC   MIN      /   WHICH SAVES MIN.
      JPAR ISZ  HRC   /GET NEXT MULTIPLYING #
      ISZ  CTM1     /OBTAINED 5 NUMBERS ?
      JMP  NXT1     /NO, GET NEXT NUMBER
      OTLP LAC   MMIN /CHECK THE MINUTES
      TCA   /        /IS THE MINUTES GREATER THAN
      TAD   MIN      /   THE MAX. NUMBER OF MINUTES
      SMA   /        /   IN AN HOUR ?
      JMP  RDERR    /YES, ASK FOR THE TIME AGAIN
      LAC   CHR      /NO, CHECK THE TOTAL TIME
      TCA   /        /IS THE TIME OF DAY #
      TAD   HR       /   LARGER THAN THE BIGGEST #
      SMA   /        /   ALLOW FOR THE TIME OF DAY ?
      JMP  RDERR    /YES, ASK FOR TIME AGAIN
      LAC   HR       /NO, PUT THE TIME INTO THE
      DAC   TIMR     /   ADDR. WHICH GIVE THE
      DAC   TIME     /   TIME OF DAY
      .TIMER 360,TOD,5 /SET UP THE TIMING R. T. SUB.
      JMS+  CONTL    /TRANSFER CONTROL TO CONTROL PROGRAM
      JMP   .+2

```

```

      .DSA      SRPO          /ADDRESS FOR CALIBRATION SURPRESSION
      .IDLE
MSGT  2000
      0
      .ASCII   "TIME"<15>
CTM1  0                /LOC. WHICH COUNTS 5 NUMBERS
CTM2  0                /USED TO IGNORE THE HOURS
SAV2  0                /LOC. TO SAVE THE MIN. #'S
HR     0                /LOC. TO SAVE THE TIME
HRC    .DSA      DECN      /POINTER FOR THE BCD MULTIPLIERS
      .DEC
DECN   10000           /BCD MULTIPLIERS
      1000
      100
      10
      1
      .OCT
CHKN  0                /SUB. TO SEPARATE OUT THE NUMBERS
      LAW      -72       /PREPARE TO LOOK AT CHARACTER
      TAD*     DT        / HEAD IN
      SAD      CODE1     /IS CHAR. A CARRIAGE RETURN ?
      JMP      RDERH     /YES, CHECK FOR POSSIBLE ERROR
      ISZ      DT        /NO,PREPARE FOR NEXT CHAR.
      SMA
      JMP      -4        /IS ASCII CHARACTER LESS THAN 72
      TAD      CODE12    /NO, GET NEXT CHARACTER
      SPA
      JMP      -7        /YES, IS CHARACTER
      JMP*     CHKN      / LARGER THAN 57 ?
                        /NO, GET NEXT CHARACTER
                        /YES, CHAR. IS A # SO EXIT
/
/##### EJECT
/#####
/
RADC  0
      JMS*     .DA       /GET VARIABLES AND PLACE ADDRESSES BELOW
      JMP      +4        /JUMP AROUND VARIABLES
NB1A  0                /ADDRESS OF THE BUFFER ADDRESS
NB2C  0                /ADDRESS OF THE WORD COUNT
NB3F  0                /ADDRESS OF THE FLAG
      DZM*     NB3F      /ZERO FLAG--WAIT FOR READ IN
      LAC*     NB1A      /INSERT BUFFER ADDRESS INTO THE
      DAC      NB5       / A/D CALL ROUTINE
      LAC*     NB2C      /INSERT THE WORD COUNT INTO THE
      DAC      NB4       / A/D CALL ROUTINE
      LAC      NB3F      /INSERT THE FLAG ADDRESS INTO THE
      DAC      NB6       / A/D CALL ROUTINE
      JMS      ADREAD    /THE A/D CALL ROUTINE (TO INIT. READ);
NB4    0                /NUMBER OF SAMPLES TO TAKE
NB5    0                /BUF. ADDRESS IN WHICH TO STORE SAMPLES
NB6    0                /COMPLETION AND ERROR FLAG ADDRESS
NB7    0                /R-T SUB. FOR INT. SERV. ROUT. TO GO TO
      JMP*     RADC      /RETURN TO FORTRAN PROGRAM
      JMS*     .DA       /SUBR. TO WRITE LIN. TABLE ON .DAT "TBI"
      JMP      +2
STA   0
REP   LAC*     STA       /ADDR. OF ADDR. OF LIN. TABLE
      DAC      STA       /SET ADDR. OF LIN.
      .INIT   TBI,OUT,TTM / TABLE
      .ENTER  TBI,STI    /PREPARE STORAGE DEVICE
      LAC      STA       /OPEN FILE FOR TABLE
      DAC      +3        /PUT ADDRESS OF TABLE INTO
      .WRITE  TBI,DUMP,0,514 / WRITE COMMAND
      .WAIT   TBI        /PUT TABLE ON STORAGE DEVICE
      .CLOSE  TBI
      JMP*     TTM
/
/##### EJECT
/#####
/
ONC4  DZM*     SUM4      /INITIALIZE TIME OF YEAR LOC.
      LAC*     TC3       /GET ADDRESS OF THE
      DAC      TC3       /
      LAC*     TC3       / DATE
      SPA
      JMP      LETMON    /GET FIRST 2 1/2 CHARACTERS
      TAD      CODE24    /ARE THEY LETTERS?
      SPA
      JMP      LETMON    /YES, CHECK FOR WORDS
      TAD      CODE24    /NO, CHECK THE NUMBERS
      SPA
      JMP      WINT      /IS THE FIRST # A ONE?
      TAD      CODE25    /YES, CHECK FOR WINTER MONTHS
      SPA
      JMP      WINT      /NO, CHECK FOR SUMMER MONTHS
      TAD      CODE25    /IS THE # >4 ?
      SPA
      JMP      ONC3     /NO, EXIT
      TAD      CODE25    /YES

```



```

SNA                /   BEEN READ IN ?
JMP                /NO, USE DEFAULT DB
LAW                /YES, ALLOW POSSIBLY ONE MORE
DAC                /   DB SETTING
ISZ                /IS ONE # STILL UNPROCESSED ?
SKP                /NO, CONTINUE THE EXIT
JMP                /YES, PROCESS THE LAST #
DAC*              /SET THE LAST LOC. TO A NEG. #
LAC                /SET THE FORTRAN COUNTER
DAC*              /   TO THE # OF DB SETTINGS
TAD                /IF THERE IS ONLY ONE DB
RAL                /   SETTING SET THE LINK
LAC                /SET UP JUMP AROUND COLL. READ IN
SZL                /IS THERE MORE THAN ONE DB SETTING ?
DAC                /NO, INSERT JUMP AROUND
JMP                /USE INPUTED DB SETTINGS
DEFDB LAC          /SET UP ASCII CODE FOR
DAC              /   ZERO DB SETTING
-----
LAC                /NO, INITIALIZE MULTIPLIER
DAC                /   TO CONVERT #'S TO MIN.
DZM                /INITIALIZE STORAGE ADDR. FOR CONVERTING
LAC                /INITIALIZE DB SETTINGS TO
DAC                /   THE BEGINING
LAC*              /SET UP COLLECTION'S DB
DAC              /   MESSAGE
CALTB .SEEK       TBI,ST1
.READ            TBI,DUMP,ST1,512      /READ TABLE IN
.WAIT           TBI
.CLOSE         TBI
LAW            -775      /PREPARES COUNTER TO CHECK
DAC            DUM2      /   THE VALIDITY OF THE LIN. TABLE
LAC            (ST1+1    /SET POINTER TO THE SECOND
DAC            DUM3      /   LOCATION OF THE TABLE
LAC            ST1      /GET THE FIRST NUMBER AND
CALLP TCA              /   COMPLIMENT
TAD*           DUM3      /IS THE PREVIOUS NUMBER
SPA              /   LARGER?
JMP            CALERA    /YES, ERRONEOUS TABLE
LAC*           DUM3      /GET PRESENT LIN. NUMBER
ISZ            DUM3      /GO TO NEXT NUMBER
ISZ            DUM2      /IS IT THE END OF THE TABLE ?
JMP            CALLP     /NO, CHECK NEXT NUMBER
DZM*           FDATA10   /YES, DELETE .DAT SLOT 10
LAS              /GET THE # 1 DATA SWICH FROM
AND            CODE40    /   CONSOLE
SZA              /IS IT SET ?
JMP            DLWT      /YES, SURPRESS PRINT OUT
.WRITE        TTO,ASC,MSG7,0
.WRITE        TTO,ASC,MSG8,0
DLWT .WRITE    TTO,ASC,MSG9,0
.WRITE        TTO,ASC,MSG10,0 /NUMBER OF HOURS
.WAIT         TTD
JMS           DBSUB     /SET UP NEXT DB SETTING
JMS           HEADM     /READ IN # OF HOURS
RUNTIM 0           /CONTAINS ADDRESS OF THE #'S READ IN
.INIT        TTD,OUT,PSTART /RESET THE #P RESTART
.INIT        TTI,IN,PSTART /   ADDRESS
LAC           CODE35    /DEFAULT VALUE FOR THE LENGTH
DAC           TC3       /   OF RUN (20 HOURS)
DZM           DSTOR     /CHECKS # OF #'S READ IN
LAW           -3        /PREPARE TO USE NO MORE
DAC           TC2       /   THAN THREE NUMBERS
GETHR  LAW       -40    /CHECK IF CHARACTER READ IN
TAD*          RUNTIM    /   IS AN ASCII # OR PERIOD
DZM*          RUNTIM    /RESET LOC. IN CHAR. BLOCK
ISZ           RUNTIM    /GO TO NEXT CHARACTER
SPA           /IS ASCII CHARACTER < 40 ?
JMP           FNUM      /YES, EXIT
SAD           CODE16    /NO, IS IT A DECIMAL POINT ?
JMP           DECP      /YES, GO TO DECIMAL POINT ROUTINE
TAD           CODE17    /NO, CHECK FOR AN ASCII NUMBER
SMA           /IS THE CHAR. < 72 ?
JMP           GETHR     /NO, GET NEXT CHAR.
TAD           CODE12    /YES, IS CHARACTER GREATER
SPA           /   THAN 57 OCTAL ?
JMP           GETHR     /NO, GET NEXT CHARACTER
JMS*          .AD       /YES, GET THE BINARY CODED DECIMAL
LAC*          MINEQC    /   EQUIVALENT OF THE NUMBER
TAD           TC1       /ADD TO THE PRECEDING NUMBERS
DAC           TC1       /   AND SAVE
ISZ           MINEQC    /SET UP NEXT MULTIPLIER
ISZ           DSTOR     /INCREMENT NUMBER COUNTER
ISZ           TC2       /HAS ENOUGH #'S BEEN OBTAINED ?
JMP           GETHR     /NO, GET NEXT NUMBER

```



```

RET3   LAW    -1          /LOAD -1 INTO MEMORY TO KEEP TRACK OF
      DAC    CNT2        / THE SAMPLE THAT HAS BEEN COLL.
      LAC    SAVEAC      /RESTORE ACCUMULATOR
END1   .RLXIT BEGIN      /RELINQUISH CONTROL TO LOWER PRIORITY
ONCE   NOP
      .TIMER 9,CHECK,5   /SET CLOCK TO WAIT 9/60 SECONDS
PK0    JMS    ADREAD     /A-D CONV. READ FOR X-SAMPLES
      TNSAM          /THE VARIABLES USED
      BUF           / ARE THE SAME ONES USED FOR
      ER2           / THE 0-SAMPLE AND ARE
      500000+RSUB    / EXPLAINED ABOVE
      LAC    SAVEAC
END2   .RLXIT BEGIN      /RELINQUISH CONTROL
RET2   LAC    TRNF       /GET DECTAPE TRANSFER FLAG
      SZA          / AND TEST IT
      JMS    DTRANS     /FLAG SET - TRANSFER DATA
A      CLAICLL        /CLEAR AC TO READ CONSOLE SWITCHES
      LAS          /GET # FROM CONSOLE'S DATA SWITCHES
      RAL          /PUT BIT 00 OF AC IN L
      SNL          /IS THE LINK A ZERO ?
      JMP    RDO        /YES, L=0 COLLECT DATA
      .TIMER 120,WT1,6 /NO, STOP COLL. AND GIVE TIME TO A
      .RLXIT BEGIN     / LOWER LEVEL BEFORE RECHECKING SWITCH
WT1    0
      DAC    SAVEAC
      LAC    MDUM1      /PUT CLOCK BACK INTO
      DAC    ONCE      / OPERATION
      DZM    CNT        /TELLS PROC. COLL. HAS STOPPED
      LAC    (JMP A     /SET UP TO RECHECK CONSOLE SWITCHES
      DAC    START     / AND PUT INTO START
      LAC    SAVEAC
      .TIMER 0,BEGIN,6 /RETURN TO R.-T. SUB. BEGIN
      .RLXIT WT1
/
/ END OF MAIN PROGRAM
/*****
BUF    .BLOCK DATBLK    /BUFFER TO STORE AD SAMPLES
NAME   .SIXBT "DATAFIDAT" /NAME OF FILE TO STORE DATA
/*****
/ INITIALIZING ROUTINE
/
INIT   LAC    (BUF1     /POINTER IN DECTAPE BUFFER
      DAC    POINT     /NAME OF DECTAPE BUFFER IN USE
      DAC    TBUF      /SAVE THE ID # (THE #
      LAC    IDCOU     / OF FRAMES PER FILE)
      DAC    CN1       /
      DZM    IDCOU     /ID NUMBER
      DZM    TRNF      /DECTAPE TRANSFER FLAG
      DZM    CNT       /COUNTER FOR CLOCK
      LAC    MDUM1     /PUT CLOCK INTO OPERATION
      DAC    ONCE
      LAW    -2        /SET BUF1 AS THE FIRST
      DAC    CIT6      / BUFFER TO BE USED
      LAC    (777      /INIT. NOISE MAX.
      DAC    MAX4      / LOCATION
      LAW    MNCC      /INITIALIZE COUNTER FOR MAXIMUM
      DAC    CIT2      / ALLOWABLE NOISE
      LAC    MCNTR     /GET THE LENGTH OF TIME REQUIRED
      TAD    TC1       / COLLECT THE PRVIOUS FILE
      DAC    DBC       /SAVE IT
      TAD    TC2       /ADD TO ALL OTHER PRVIOUS TIMES TO
      DAC    TC2       / COLLECT THE OTHER FILES, SAVE AND
      TAD    TC3       / COMPARE TO THE MAXIMUM TIME
      SMA          /ARE THE FILE TIMES LARGER ?
      JMP    EXTIP     /YES, STOP COLLECTION
/
/ DECTAPE FILE ROUTINE
/ STORE DATA IN FILE ACCORDING TO RESPONSE
/
UP     .FSTAT OUTPT,NAME /CHECK FOR COLLECTED FILE
UPI    SKP          /(<REPLACED BY "SZA">)IS FILE PRESENT?
WRITE  LAC    UPDATE  /YES, ACKNOWLEDGE THE PRESENCE
      TCA    MCNTR    /GET THE BEGINNING TIME FROM
      DAC    TC1      / THE BINARY CLOCK COUNTER
      .DELETE OUTPT,NAME / TO DETERMINE THE COLL. TIME
      .ENTER  OUTPT,NAME /DELETE FILE IF PRESENT
C3     .ENTER  OUTPT,NAME /OPEN FILE
      JMS    DTRANS   /WRITE DUMMY BLOCK
      JMS    DTRANS   /TWICE
      LAC    MDUM2    /TELLS PROCESSING THAT COLLECTION
      DAC    REPL     / HAS STARTED COLLECTING A FILE
      JMP    RDO      /RETURN

```



```

/
/#####
/#####
/
/ ROUTINE TO CLOSE FILE AND SET UP PARAMETERS FOR
/ COLLECTION AND PROCESSING
/
RESTAR LAC TRANF /COMPLETE TRANSFER IF NECESSARY
MDUM2 SZA
JMS DTRANS
C6 .WRITE OUTPT,DUMP,SEVN,2 /WRITE END OF FILE ID
C2 .WAIT OUTPT
C7 .CLOSE OUTPT
LAC MDUM1 /INITIALIZE THE LOCATIONS
DAC START / "START" FOR COLLECTION
DAC REPL / AND "REPL" FOR PROCESSING
LAC SUM4 /GIVE THE SUMMATION OF THE LOWEST
DAC* SUM5 / NOISE TO THE PROCESSING
DZM* CHG /RESET DB CHANGER
LAC MDUM2 /IGNORE THE FIRST
DAC UPI / FILE ONLY
JMP START /GO TO START
/
/#####
/#####
/
/ CLOCK INTERRUPT ROUTINE FOR AUTOMATIC O-SAMPLE START
/
CHECK 0
DAC SAV /SAVE AC
DZM BEGIN /ZERO R-T SUB. TO AVOID POSS. ERROR
LAC CNT /EXAMINE COUNTER
RAH /LOWEST BIT OF CNT IN L
SZL
JMP BKUP /L=1-CNT ODD-HALF FRAME DURING 9/60 SEC
JMP NORM /L=0-CNT EVEN-FULL FRAME DURING 9/60 SEC
BKUP DZM CNT /CLEAR HALF FRAME THAT WAS TAKEN:
RAH /RESTOR LINK
LAW -1 /RESET IDCOU BACK ONE FRAME
TAD IDCOU
DAC IDCOU
LAW -1 /RESET BUFFER COUNTER
TAD CTT8 / BACK ONE
DAC CTT8 / FRAME
LAW -DPR /RESET POINT BACK HALF FRAME
TAD POINT
DAC POINT
LAC MDUM1 /PUT CLOCK BACK INTO OPERATION
JMP EXT1 /PREPARE TO EXIT
NORM RAH /RESTORE THE LINK
EXT1 LAC (JMP PKO /STOP CLOCK FROM OPERATION
DAC ONCE
LAC (JMP RDO /PREPARE TO COLLECT
DAC START / AN O-FRAME
LAC SAV /RESTORE AC
TIMER 0,BEGIN,6 /RETURN TO POINT AT WHICH INTERRUPT
RLXIT CHECK / OCCURED
/
/#####
/#####
/
/ SUBROUTINE TO DO A TABLE LOOKUP FOR DATA LINEARIZATION.
/ ALSO SAVES SUM AND MAXIMUM OF EACH SET OF 5 NOISE SAMPLES,
/ SETS NEGATIVE NUMBERS TO ZERO, AND IF NECESSARY JUMPS AROUND
/ "SKAR" NUMBER OF DATA NUMBERS BETWEEN THE NOISE AND DATA POINTS.
/
LIN 0
ISZ BPOINT /DATA STARTS AT BJT
LAC* BPOINT /GET INPJT DATA WORD
AND (1777 /MASK ANY EXTRA BITS
TAD (-1000 /CHECK FOR NEG. #'S
SMA
JMP ERI /NEG # FOUND
TAD (ST1+1000 /LOCATE # IN TABLE
DAC DJM1 /SET ADDRESS OF NUMBER
LAC* DUM1 /LOAD LINEARIZED # INTO AC
SHR
ERI CLARCLL /SET NEG # TO ZERO
DAC DUM1 /STORE LINEARIZED #
ISZ CTT5 /IS THIS THE FIFTH NOISE SAMPLE ?
JMP +5 /NO, SKIP AROUND CODE
LAC (SKAR /YES, SKIP AROUND "SKAR" DATA
TAD BPOINT / OF THE FRAME BEING
DAC BPOINT / COLLECTED
LAC DJM1 /RESTORE THE LIN. DATA #
ISZ CTT1 /IS IT A NOISE SAMPLE ?
JMP DONOS /YES, PROCESS NOISE SAMPLE
LAW -1 /RESET THE TWO COUNTERS
DAC CTT1 / FOR THE DATA OF THE

```



```

/
ADWCR=26 /A-D WORD COUNT
ADCAR=ADWCR+1 /AND CURRENT ADDRESS REGISTERS
.SCOM=100 /MONITOR'S COMMUNICATION AREA
ADWI=703724 /A-D CONVERTER WRITE INITIALIZE
ADSO=703701 /SKIP ON WORD COUNT OVERFLOW
ADST=703721 /SKIP ON DATA TIMING ERROR
ADCO=703704 /CLEAR OVERFLOW FLAG
ADCT=703744 /CLEAR TIMING FLAG
/
/ ENTRY POINT FOR A-D INTERFACE INITIALIZATION
/
ADREAD 0

JMP INSET /REPLACED BY "LAC* ADREAD"
TCA
DAC* (ADWCR) /SET WORD COUNT
ISZ ADREAD
LAW -1
TAD* ADREAD /BUFFER ADDRESS -1
DAC* (ADCAR) / TO CURRENT ADDRESS REG.
ISZ ADREAD
LAC* ADREAD /GET FLAG ADDRESS
DAC INFLAG
DZM* INFLAG /CLEAR FLAG
ISZ ADREAD
LAC* ADREAD /GET REAL-TIME SUBROUTINE ADDRESS
DAC INSUB
ISZ ADREAD /POINT TO RETURN LOCATION
ADWI /INITIALIZE INTERFACE
JMP* ADREAD /RETURN
/
/ THE FOLLOWING CODE IS EXECUTED ONLY ONCE
INSET LAC* (.SCOM+55) /GET ENTRY POINT ADDRESS OF .SETUP
ADSV A DAC .
SAV LAC* (.SCOM+51) /ENTRY POINT OF REALTP
REALTP DAC .
LAC (400010) /RAISE THE API
ISA / LEVEL
JMS* ADSVA /CALL .SETUP TO CONNECT
ADSO / ADINT TO
ADINT / THE API
DBK /DEB BREAK FROM API LEVEL
LAC (LAC* ADREAD)
DAC ADREAD+1 /MODIFY INSTRUCTION
JMP ADREAD+1 / AND JUMP TO IT
/
/
/ INTERRUPT SERVICE ROUTINE. EXECUTED IMMEDIATELY AFTER COMPLETION
/ OF DATA TRANSFER. DETERMINES STATUS OF A-D INTERFACE, SETS
/ COMPLETION FLAG AND ACTIVATES REAL-TIME SUBROUTINE.
/ RUNS AT API LEVEL 0.
/
/
ADINT 0
DBA /PAGE ADDRESSING MODE
DAC ADSVA /SAVE AC
ADST /TIMING ERROR?
SKP!CLAIAC /NO,+1 TO AC
LAW -1001 /YES, ERROR CODE
DAC* INFLAG /SET FLAG
ADCO /CLEAR
ADCT / INTERFACE FLAGS
LAC* (.SCOM+102) /RAISE TO API
ISA / LEVEL 0 OR 1
LAC INSUB /REAL-TIME SUBROUTINE ADDRESS
SNA
JMP ADXIT /BYPASS MONITOR CALLS IF ZERO

ADXIT JMS* REALTP /ACTIVATE REAL-TIME SUBROUTINE
LAC (404000) /REQUEST AN API INTERRUPT
ISA / AT SOFTWARE LEVEL 4
LAC ADSVA /RESTORE AC
DBR /SET TO LEAVE HARDWARE API LEVEL
JMP* ADINT /RETURN TO INTERRUPTED PROGRAM
/
/*****
.EJECT
/XXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXX
/ PROCESSING'S MACRO PROGRAMS. THEY INITIALIZE THE SOTRAGE
/ DEVICE, WAIT FOR FILE TO BE COLLECTED, CHECK FOR UNWANTED
/ COLLECTION STOPAGE, SWITCH STORAGE DEVICES IF NECESSARY, TELL
/ WHEN THE PROGRAMS HAVE REACHED THE END OF RUN, GIVES TIME TO
/ BACKGROUND, AND READS IN DATA AND UNPACKS IT.

```

```

CHNG      0
          JMS*   .DA           /LOAD PARAM. ADDR. IN SUM5 AND CHG
          JMP     .+4           /SKIP OVER PARAM. LIST

SUM5      0
CHG       0
CMULC    0
          ISZ    CTT10        /LOC. TO CHANGE MUL. CONSTANT
          SKP    /RESTART THE DEVICE ?
          JMP     AGN          /NO
          ISZ    FIL+1        /YES, RESET PARAMETERS TO BEGIN. OF DEV.
          CONTR /CHECK IF COLLECTION IS
          LAC    FIL+1        / FINISHED COLLECTING NEW FILE
          CMA    /IS THE COLLECTION
          TAD    NAME+1       / RECOLLECTING
          SPA    / THE FIRST FILE
          LAW    -1           / AGAIN ?
          DAC    CTT10        /NO, SET THE COUNTER TO JUMP AROUND
          LAW    -1           /YES, SET COUNTER TO A POS. #
          DAC    CTT9         /PREPARE TO READ
          LAC    BUF31        / DATA
          DAC    POINT2       /RESET BUF3 POINTER WITH
          LAC    (ISZ SWITC   / THE ADDR. OF BUF3
          DAC    LCA          /PREPARE TO READ TWO
          C13     .INIT        / DUMMY BLOCKS
          RT      .FSTAT DATIN,FIL /CHECK FOR PRESENTS OF NEW FILE
          JMP     ERR3        /FILE NOT PRESENT
          C11     .SEEK DATIN,FIL /PREPARE TO READ
          JMP*    CHNG        /RETURN TO FORTRAN PROGRAM
          FIL     .SIXBT "DATA#DAT" /NAME OF DATA FILE USED BY PROC.
          ERR3   .WRITE TTO,ASC,MSG6,0 /FILE NOT FOUND
          MSG6   JMP     CMULC+1 /ERASED FILE ? LOOK FOR NEXT FILE
          2000
          0
          .ASCII "FILE NOT FOUND"<15>
          AGN    LAC    AFI          /INITIALIZE DATA-FILE
          DAC    FIL+1         / NAME
          NDK3   SKP    /DETERMINES # OF STORAGE DEVICES
          JMS    SWIT          /CHANGE STORAGE DEVICES
          JMP     DUM7+1       /NO, CONTINUE PROCESSING
          ENDPK  .WRITE TTO,ASC,END,0 /END OF PROC. MESSAGE
          .WAIT TTO
          ANSE   JMS    READM        /WAIT FOR RESPONSE AND PUT
          0       / THE ADDR. OF IT HERE
          LAC*   ANSE
          SAD    (120         /IS THE RESPONSE A "P" ?
          JMP     RSTR        /YES, SET UP TO RESTART EVERYTHING
          .IDLE /GIVE COMPLETE CONTROL TO BG.
          RSTR   .CLEAR DATIN     /CLEAR COLLECTION DEVICE
          .WRITE TTO,ASC,MSG2,0 /DB MESSAGE
          .WAIT TTO
          JMS    DBSUB        /SET UP NEXT DB SETTING
          JMS    READM        /WAIT FOR REPLY
          DMRPLY 0            /THE REPLY
          DZM    MCNTR        /ZERO BINARY TIMER
          .TIMER 0,BEGIN,6    /RESTART COLLECTION
          JMP     AIN         /RESTART PROC.
          / SUBROUTINE USED TO WAIT FOR FILE TO BE COLLECTED AND STORED
          / WHILE GIVING TIME TO BACKGROUND
          CONTR 0
          WT2   LAC    NAME+1     /COLLECTION'S FILE NAME
          SAD    FIL+1         /IS PROC.'S FILE NAME THE SAME ?
          MDUM4 SKP    /YES, GO TO WAIT
          JMP*   CONTR        /NO, RETURN TO PROCESS FILE
          ISZ    CTT11        /END OF THE
          SKP    / RUN ?
          JMP     ENDPK       /YES,EXIT
          .TIMER 30,WAIT1,3    /YES, RELINQUISH TIME TO BG.
          LAC    SAVAC        /RESTORE AC
          .IDLE /WAIT FOR CLOCK INTERRUPT
          /REAL TIME SUB.--USED TO ALLOW TIME FOR BACKGROUND
          WAIT1 0
          DAC    SAVAC        /SAVE AC
          DZM    WAIT1        /ZERO R-T SUB. ENTRY PT. TO ALLOW REENTR
          JMS    CKCOL        /CHECK FOR COLLECTION STOPPAGE
          JMP     #I2         /CHECK FOR END OF COLL.
          / SUBROUTINE TO CHECK FOR UNWANTED COLL. STOPPAGE
          CKCOL 0
          LAY    -2
          TAD    CNT          /HAS COLL. ENDED ALL COLLECTING
          SPA    / FOR TODAY?
          JMP*   CKCOL        /YES, RETURN
          LAC    TST1        /NO, HAS COLLECTION STOPPED

```



```

      TGA          / READING
      TAD          IDCOU / IN DATA?
REPL  SZA          / (REPLACED BY "NOP" WHEN COLL. IS DONE)
      JMP          SET1 /NO, RESET TESTER AND RETURN
      LAC          MDUM1 /YES, FREE TIMER
      DAC          ONCE / OPTION IN COLL.
      ISZ          CTT7 /TELLS IF STOPPAGE OCCURED
      .WRITE      TTO,ASC,STPM,0 /RING BELL
      .TIMER      0,CHECK,5 /RESTART COLLECTION
SET1  LAC          IDCOU /SET ID # INTO
      DAC          TST1 / TESTER FOR STOPPAGE
      JMP*         CKCOL /CHECK FOR COLL. TO BE FINISH
/ SUB. TO CHANGE THE .DAT SLOT #'S TO CHANGE STORAGE DEVICES
SWIT  0
      LAC          (DATIN /SWITCH STORAGE DEVICES
      ISZ          CTT4 / BY CHANGING THE
      LAC          (DATIN2 / .DAT SLOT IN COMMANDS:
      DAC          C11 / .SEEK
      DAC          C13 / .INIT
      DAC          C14 / .WAIT
      DAC          C15 / .CLOSE
      DAC          RSTR / .CLEAR
      TAD          (3000
      DAC          RT / .FSTAT
      TAD          (1000
      DAC          LBA / .READ
      LAC          CTT4 /IS DEVICE ON .DAT SLOT "DATIN"
      SMA          / TO BE PROCESSED?
      LAW          -2 /YES, RESET DEVICE CONTROLLER
      DAC          CTT4 /MAKE ANY CHANGE IN CONTROLLER
      JMP*         SWIT
/
PP7  0
      LAS          /GET THE CONSOLE DATA SWITCH
      AND          (40000 / NUMBER 3
      SNA          /IS IT A 1 ?
      JMS*         WAPP /NO, PRINT DATA OUT ON PAPER TAPE
      JMP*         PP7 /RETURN TO PROC
/
/XXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXX
. EJECT
/XXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXX
/ .READ, DUMP MODE FROM DECTAPE ON A VARIABLE .DAT SLOT
/ FILLS 252 DEC WORD BUFFER AND OUTPUTS 26
/ WORDS TO ARRAY IDAT EVERY TIME CALLED.
/ THESE ARE UNPACKED FROM 18 WORDS OF THE BUFFER.
/ IDAT: WORD 1 I.D. #
/ WORD 2-6 NOISE SAMPLES
/ WORD 7-27 DATA
/ NEGF: SET IF A NEGATIVE NUMBER WAS IN THE DATA
/
DUMPT 0
      JMS*         .DA /PICKUP ADDR OF ADDR
      JMP          .+3 /OF ARRAY
A2  0
FLAG  0
      LAC*         A2 /SET ON NEG #
      DAC          A2 /GET ADDR.
      LAW          -NDPC / OF ARRAY
      DAC          COUNT /SET COUNTER OF DATA TO BE
      ISZ          CTT9 / PROCESSED
      JMP          LBB /GET POINTER
      LAW          -DATSTR /NO, CONTINUE WITH PRESENT SET OF DATA
      DAC          CTT9
LBA .READ DATIN,DUMP,BUF3,DTBLK /GET 1 BLOCK OF DATA
C14 .WAIT DATIN
LCA ISZ SWITC /INITIALLY READ TWO DUMMY BLOCKS
      JMP          LBA
      LAW          -3 /RESET CONTROL TO READ
      DAC          SWITC / TWO DUMMY BLOCKS
      LAC          (JMP LCB /READ ONE BLOCK OF DATA
      DAC          LCA / AT A TIME
LCB LAC BUF31 /GET ADDRESS OF BUF3
      DAC          POINT2 / POINT2 TO BUF3
LBB LAC* POINT2 /GET THE ID (FIRST WORD IN DATA SET)
      SAD          SEVN /END OF FILE ID?
      JMP          ENF /YES, RESET PARAM'S AND CLOSE FILE
      LAC*         POINT2 /GET ID AND PUT
      DAC*         A2 / INTO THE FORTRAN ARRAY
      ISZ          A2 /GO TO NEXT ADDR. OF THE ARRAY
LOOP ISZ POINT2 /GO TO NEXT DATA WORD
      LAW          -2 /PREPARE TO UNPAC
      DAC          TC10 / TWO DATA WORDS
      LAC*         POINT2 /GET DATA WORDS FROM BUF3

```



```

C
  SUBROUTINE TBFORL
    INTEGER START(512),DLPO,STAT,IAS(44)
    REAL KI
    COMMON /TA/ S(43),TU(44),TUO(44)
    DATA C,CS,CN,NR/1HC,1HS,1HN,120784/
C   DETERMINE IF CALIBRATION IS NEEDED AND WHICH PRINTOUT TO USE
7   WRITE(6,105)NR
105  FORMAT(18H WHICH CALIBRATION/24H (S-SHORT, R-REGULAR, OR
    137H C-COMplete PRINT OUT OR N-NO CALIB.),A2)
    READ(4,204)CAL
204  FORMAT(A1)
    IERR=0
    DLPO=0
C   SET UP THE WANTED CALIBRATION
    IF(CAL.EQ.CN)RETURN
    IF(CAL.EQ.CS)DLPO=-1
    IF(CAL.EQ.C)DLPO=1
C   NUMBER OF A/D NUMBERS TO READ PER DB SETTING=NAVI*NAV
    NAVI=5
    NAV=500
    AV=NAVI*NAV
    ICT=0
C   NUMBER OF ATTENUATOR SETTINGS=NT1
    NT1=42
    NT2=NT1-1
    NT3=NT1-2
C   INPUT SIGNALS--FROM 5DB (OF 1000) TO INFINITY
    TU(NT1)=562.34
    TU(NT2)=501.19
    TU(NT3)=446.68
    TU(NT1-3)=398.11
    TU(NT1-4)=354.82
    TU(NT1-5)=316.23
    TU(NT1-6)=281.84
    TU(NT1-7)=251.19
    TU(NT1-8)=223.87
    TU(NT1-9)=199.53
    TU(NT1-10)=177.83
    TU(NT1-11)=158.49
    TU(NT1-12)=141.25
    TU(NT1-13)=125.90
    TU(NT1-14)=112.20
    TU(NT1-15)=100.00
    TU(NT1-16)= 89.13
    TU(NT1-17)= 79.43
    TU(NT1-18)= 70.80
    TU(NT1-19)= 63.10
    TU(NT1-20)= 56.24
    TU(NT1-21)= 50.12
    TU(NT1-22)= 44.67
    TU(NT1-23)= 39.81
    TU(NT1-24)= 35.48
    TU(NT1-25)= 31.62
    TU(NT1-26)= 28.18
    TU(NT1-27)= 25.12
    TU(NT1-28)= 22.39
    TU(NT1-29)= 19.95
    TU(NT1-30)= 17.78
    TU(NT1-31)= 15.85
    TU(NT1-32)= 14.13
    TU(NT1-33)= 12.59
    TU(NT1-34)= 11.22
    TU(NT1-35)= 10.00
    TU(NT1-36)= 8.91
    TU(NT1-37)= 7.94
    TU(NT1-38)= 7.08
    TU(NT1-39)= 6.31
    TU(NT1-40)= 5.62
    TU(1)= 0.00
    TUO(NT1)=512.
C   SET UP MESSAGES FOR TELLING WHICH ATTENUATOR SETTING TO DO
    DO 11 I=1,NT3
11   IAS(I)=I+5
    IAS(NT2)=99
C   DO LOOP TO INPUT ALL THE OUTPUTED SIGNALS
    DO 2 I1=1,NT2
    I2=NT1-I1
    DUM=0
    WRITE (6,100) IAS(I1),NR
100  FORMAT(8H SET TO ,I2,23HDB ATTENUATION AND C.R.,A2)
    READ(4,200)F
200  FORMAT(F5.1)
C   IF THE INPUTED NUMBER IS TWO DIGITS RESTART THE SETTINGS
    IF(F.GT.10.)GO TO 23

```

```

C   ISSUE AN A/D CONVERTER READ "NAVI" TIMES
DO 10 J1=1,NAVI
C   READ "NAV" NUMBERS FROM THE A/D CONVERTER
CALL RADC(START,NAV,STAT)
400  IF(STAT.EQ.0)GO TO 400
DO 1 J=1,NAV
IF(START(J).GT.511) START(J)=0
C   STORE THE INPUTED NUMBERS IN A REAL VARIABLE
1   DUM= DUM+FLOAT(START(J))
10  CONTINUE
C   AVERAGE THE OUTPUTED NUMBERS AND GO TO THE NEXT SETTING
2   TUO(I2)=DUM/AV
DO 3 I1=1,NT2
I2=I1+1
ICT=ICT+1
I3=NT1-I1
C   SET UP THE SLOPES OF EACH LINE SEGMENT APPROXIMATION
S(I1)=(TU(I2)-TU(I1))/(TUO(I2)-TUO(I1))
IF(ICT.EQ.2)GO TO 3
C   WRITE OUT STRAIGHT LINE APPROXIMATION TABLE
WRITE(6,102) I1,S(I1),I1,TU(I1),I1,TUO(I1),IAS(I3)
ICT=1
IF(DLPO.GT.-1)ICT=0
C   POSSIBLE ERROR CONDITIONS FOR THE APPROXIMATION JUST FORMED
3   IF(S(I1).LT..3.OR.S(I1).GT.7.) IERR=IERR+1
102  FORMAT (4H S(,I2,2H)=,F6.3,5X,3HTU(,I2,2H)=,F8.3,5X,
14HTUO(,I2,2H)=,F8.3,3X,I2,2HDB)
C   WRITE OUT LAST VALUES OF THE TABLE
WRITE(6,101) TU(NT1),TUO(NT1)
101  FORMAT(19X,7HTU(42)=,F8.3,5X,8HTUO(42)=,F8.3,7H 5DB///)
IERR2=0
C   FINAL OUTPUT VALUE FOR THE LINEARIZATION TABLE FORMATION
XF=511.5
C   GET THE INPUT VALUE AND STORE IN "XF"
CALL LINAP(XF,NT1)
IF (XF.LE.0.)XF=512
C   NORMALIZATION FACTOR OF THE OUTPUT VALUES OF THE LIN. TABLE
X1=1023./XF
C   INITIAL OUTPUT USED TO DETERMINE THE INPUT VALUE
X1=511.5/1024.
C   FIRST INPUT VALUE OF THE TABLE
START (1) =1
DO 4 I=2,512
C   NEXT OUT PUT VALUE USED
X3=X1*FLOAT(2*I-1)
C   GET THE INPUT VALUE AND STORE IN "X3"
CALL LINAP (X3,NT1)
X3=(X3*X1+1.)/.2.
IF(X3.LE.1.)X3=1.1
C   STORE INPUT VALUE IN INTEGER LIN. TABLE
START(I)=IFIX(X3)
C   ERROR CONDITION FOR LINEARIZATION TABLE
4   IF(START(I)+1.LT.START(I-1))IERR2=IERR2+1
C   LAST VALUE OF THE TABLE
START(512)=511
IF(DLPO.LT.1) GO TO 6
C   NEW PAGE
WRITE (6,103)
103  FORMAT (1H1)
C   WRITE LINEARIZATION TABLE ON TELETYPE
WRITE (6,104) (START(I),I=1,512)
104  FORMAT (10(15,2X))
C   NEW PAGE
WRITE(6,103)
C   WRITE TABLE ON A STORAGE DEVIVE (DUMP MODE)
6   CALL TIM(START)
C   WRITE OUT ANY ERROR AND ALLOW RECALIBRATION IF NEEDED
IF(IERR2.NE.0)GO TO 9
IF(IERR.EQ.0)RETURN
WRITE(6,106)
106  FORMAT(//47H ****CHECK CALIBRATION FOR POSSIBLE ERRORS****//)
GO TO 7
9   WRITE(6,107)
107  FORMAT(37H +++ ERROR--BAD CALIBRATION TABLE +++
GO TO 7
RETURN
END

```

```

C *****SUBROUTINE LINAP*****
C LINAP TRANSFORMS OUTPUT VOLTAGES INTO INPUT VOLTAGES OF
C THE RECEIVER. THE CALIBRATION DATA IS CONTAINED IN SUB-
C ROUTINE VALUE.
C *****
C INPUT AND OUTPUT:
C A IS THE OUTPUT VOLTAGE THAT IS TRANSFORMED INTO
C INPUT VOLTAGE
C NCS IS THE NUMBER OF DB SETTINGS
C
C SUBROUTINE LINAP(A,NC5)
C COMMON /TA/ S(43),TU(44),TUO(44)
C N=1
C DIVIDE THE STRAIGHT LINE APPROXIMATION INTO 4 AREAS
C NC54=(NC5+2)/4
C NC53=NC5*3/4
C NC52=(NC5+1)/2
C FIND WHERE THE INPUTED NUMBER LIES
C IF(A.GT.TUO(NC54))N=NC54
C IF(A.GT.TUO(NC52))N=NC52
C IF(A.GT.TUO(NC53))N=NC53
C SET THE UPPER LIMIT OF THE SEARCH
C K=N+NC54
C SEARCH FOR THE CORRECT LINE SEGMENT
C DO 5 I=N,K
C J=I+1
C IF(A.GT.TUO(I).AND.A.LE.TUO(J)) GO TO 10
S CONTINUE
C LINE SEGMENT COULDN'T BE FOUND
C A=0.
C RETURN
C GET THE VALUE OF THE CORRESPONDING OUTPUT
10 A=(A-TUO(I))*S(I)+TU(I)
C RETURN
C END

```

```

C *****PROC*****
C THIS PROGRAM OPERATES ALONG WITH DLOGF1. IT TAKES DATA READ
C BY A SUBROUTINE AND CHECKS THE NOISE AND SATURATIONS
C FOR GOOD DATA. IT SUMS THE SQUARES OF THE DATA AT EACH HEIGHT
C AND THE ACCEPTABLE NOISE, SUBTRACTS THE ACCEPTABLE NOISE SQUARED
C FROM THDATA, TAKES THE SQUARE ROOT OF THE DATA PRESERVING THE
C SIGN, TAKES THE SQUARE ROOT OF THE NOISE, AND PRINTS OUT THE RESULTS
C ON THE TELETYPE. THE RESULTS MAY ALSO BE PRINTED OUT ON PAPER TAPE
C IF NEEDED. THE PROGRAMS USED BY PROC ARE:
C CHNS----INITIALIZES STORAGE DEVICE AND WAITS FOR FILE
C TO BE COLLECTED PLUS COMMUNICATING WITH COLL.
C DRD73--STORES DATA IN REAL ARRAY (FORTRAN)
C CALC2---CALCULATES ELECTRON DENSITIES (FORTRAN)
C CKCOL---CHECKS FOR UNWANTED STOPS IN COLLECTION
C PROGRAM (MACRO)
C PP7-----CHECKS DATA SWITCHES ON THE CONSOLE TO ALLOW
C OR DISALLOW RESULTS TO BE PUT ON PAPER TAPE
C *****
C SUBROUTINE PROC(IA,PLF,ITIME)
C COMMON /PPC/ AO(21),AX(21),AVAO(21),AVAX(21),ITIM(4),
C IXO(21),IRJ(21),BNO(5),BNX(5),EL(21),
C 2RBMX,AVNO,BMO,AVNX,BMX,ID,IR
C COMMON /STAT/ IDB(4),RDATE(2),REAS(5),IDBC,BMXNS,NC4
C SRBMX=3.1
C RBMX=SRBMX*SRBMX
C IWIN=-1
C IEQU=0
C ISUM=1
C NC4=1
C INITIALIZE PAPER PUNCH
C WRITE(7,1500)
C INITIALIZE ALL VARIABLES NEEDED TO BE INIT.
10 SNO=0.
C SNX=0.
C IR=0
C IRN=0
C BMO=0.
C BMX=0.
C DO 100 I=1,21
C AVAO(I)=0.
C AVAX(I)=0.
C AO(I)=0.
C AX(I)=0.
100 IRJ(I)=0

```

```

DO 110 I=1,4
  BNO(I)=0.
  BNK(I)=0.
110  C INITIALIZE STORAGE DEVICE AND
  C PREPARE TO READ DATA
      CALL CHNG(IBM,IBMX,IBCH,IMCC)
      IF(IMCC.LE.0.OR.IMCC.GT.9)GO TO 20
      CMC=FLOAT(IMCC)-4.5
      RBMX=RBMX+CMC*CMC*.4
20   IF(IBM.LT.1)IBMX=216
      IMCC=0
  C NOISE CRITERION
  C IBMX=THE SUM OF THE SET OF 45 NOISE SAMPLES WHICH HAS THE
  C MIN. MAXIMUM OF ALL THE MAXIMUMS OF EACH SET OF 45 NOISE SAMPLES
  C THE AVERAGE NOISE FOR THE FIRST 5 IN EACH FRAME HAS
  C TO BE LESS THAN SRBMX*IBMX/45. WHERE RBMX IS THE SUPPLIED CONSTANT.
  C FOR SPEED RBMX*(IBMX/45)**2*5. IS COMPARED TO THE SUM OF THE
  C SQUARES OF THE NOISE.
      DUMX=FLOAT(IBM)/45.
      IF(DUMX*SRBMX.GT.500.)DUMX=500./SRBMX
      BMXNS=RBMX*DUMX*5.*DUMX
      KEFO=0
      KEAFX=0
      ID=0
  C GET ORDINARY MODE DATA
  C 30 CALL DRD73(AO,BNO,IEHR,ID,KEFO)
      IF(KEFO.EQ.1)GO TO 80
  C GET EXTRAORDINARY MODE DATA
      CALL DRD73(AX,BNK,IEHR,ID,KEAFX)
      IF(KEAFX.EQ.1)GO TO 80
  C SET UP CHECK FOR REJECTION BECAUSE OF NOISE CRITERION
      BMEANO=0.
      BMEANX=0.
      DO 120 I=1,5
120  BMEANO=BMEANO+BNO(I)*BNO(I)
      BMEANX=BMEANX+BNK(I)*BNK(I)
      IF(BMEANO.GT.BMXNS.OR.BMEANX.GT.BMXNS)GO TO 50
  C NOISE USED TO SUBTRACT FROM DATA SAMPLES
      BMO=BMO+BMEANO
      BMX=BMX+BMEANX
  C SUM OF THE SQUARE OF THE UNSATURATED DATA AT EACH HEIGHT
      DO 140 I=1,21
35   IF(AO(I).GE.510..OR.AX(I).GE.510.)GO TO 40
      AVAO(I)=AVAO(I)+AO(I)*AO(I)
      AVAX(I)=AVAX(I)+AX(I)*AX(I)
      GO TO 140
  C REJECTIONS DUE TO SATURATIONS OF DATA
40   IRJ(I)=IRJ(I)+1
140  CONTINUE
      GO TO 60
  C REJECTIONS FROM NOISE CRITERION
50   IR=IR+1
  C SET UP AVERAGE NOISE USED IN REJECTION CRITERION
60   SNO=BMEANO+SNO
      SNX=BMEANX+SNX
      GO TO 30
80   ID=ID/2
      BID=ID*5
  C MAXIMUM ALLOWABLE NOISE
      BMXNS=BMXNS/(DUMX*5.*SRBMX)
  C RMS OF ALL NOISE SAMPLES
      AVNO=SQRT(SNO/BID)
      AVNX=SQRT(SNX/BID)
  C NUMBER OF ACCEPTABLE NOISE SAMPLES
      RN=5*(ID-IR)
      DO 150 I=1,21
  C NUMBER OF REJECTIONS AT EACH HEIGHT
      IRJ(I)=IRJ(I)+IR
  C NUMBER OF ACCEPTABLE DATA AT EACH HEIGHT
      RSAM=ID-IRJ(I)
  C AVERAGE SUM SQUARED OF ACCEPTABLE DATA FOR EACH HEIGHT MINUS THE
  C AVERAGE SUM SQUARED OF THE ACCEPTABLE NOISE
      AVOC=AVAO(I)/RSAM-BMO/RN
      AVXC=AVAX(I)/RSAM-BMX/RN
  C THE RMS OF THE ACCEPTABLE DATA AT EACH HEIGHT (PRESERVING THE SIGN)
      AVAO(I)=(ABS(AVOC)/AVOC)*SQRT(ABS(AVOC))
      AVAX(I)=(ABS(AVXC)/AVXC)*SQRT(ABS(AVXC))
      EL(I)=0.
      XO(I)=0.
      IF(AVAO(I).LE.0.OR.AVAX(I).LE.0)GO TO 150
      XO(I)=AVAX(I)/AVAO(I)

```

```

150 CONTINUE
C THE RMS OF THE ACCEPTABLE NOISE
  BMO=SQRT(BMO/RN)
  BMX=SQRT(BMX/RN)
  CALL CALC2(XO,1.20,EL,IA)
C GET THE TIME OF DAY
  DO 155 I=1.4
  I1=5-I
155 ITIM(I)=ITIME/(10**I1)
C WRITE THE HEADING ON THE TELETYPE
  WRITE(6,1050)ITIM,RDATE,REAS,IDB(NC4)
1050 FORMAT(1H1.4I1.4X,2A5.3X,5A5.3X,I2.2HDB)
  WRITE(6,1100)BMXNS,RBMX
1100 FORMAT(//19H MAX. ALLOW. NOISE=F7.1,16H MULT. CONST.=F7.3/)
  WRITE(6,1200)AVNO,BMO,AVNX,BMX
1200 FORMAT(16H O-NOISE AV.(1),F8.1,7H (2),F8.1/
  116H X-NOISE AV.(1),F8.1,7H (2),F8.1)
  WRITE(6,1300)ID,IR
1300 FORMAT(//1X,I4,8H SAMPLES,5X,I5,12H REJ.(NOISE)//3X,
  14HREJ./48H (N.+SAT.) HEIGHT AV. AO AV. AX AX/AO ED/)
  HT=5 8.5
  DO 160 I=1,21
C CHECKS FOR COLLECTION STOPPAGE
  CALL CKCOL
  HT=HT+1.5
  WRITE(6,1400)IRJ(1),HT,AVAO(1),AVAX(1),XO(1),PLF,EL(1)
1400 FORMAT(4X,I4.4X,F5.1,3X,F6.1,2X,F6.1,2X,F5.2,A3,F6.0)
160 CONTINUE
C ALLOWS RESULTS TO BE SAVED ON PAPER TAPE
  CALL PP7
C NEXT ATTENUATOR SETTING
  NC4=NC4+1
  IF(NC4.3T.IDBC)NC4=1
  IF(IDBCH.GT.0)IDB(NC4)=5*IDBCH
  GO TO 10
1500 FORMAT(1H )
  RETURN
  END

C*****SUBROUTINE DRD73*****
C DREAD READS 21 SAMPLES OF SIGNAL AND 5 SAMPLES OF NOISE
C FROM DECTAPE. THE OUTPUT VOLTAGES HAVE BEEN TRANSFORMED INTO
C INPUT VOLTAGES. THE PROGRAM USES SUBROUTINE DUMPT (MACRO) TO
C READ DATA FROM STORAGE DEVICE.
C*****
C
  SUBROUTINE DRD73(A,BMEAN,IERR,ID,KEOF)
  DIMENSION A(21),IDAT(27),BMEAN(5)
  KEOF=0
  N=5
  N1=N+1
  N2=N+2
  N3=N1+21
C GET ONE SET OF DATA (26 NUMBERS)
  CALL DUMPT(IDAT,NESF)
C CHECK ID CONSECUTIVE
  IF(ID-IDAT(1)+1) 10,15,10
C CHECK FOR END OF FILE
10 IF(IDAT(1).NE.130050) 30 TO 20
C TELL PROC IT'S THE END OF THE FILE
  KEOF=1
  RETURN
15 ID=IDAT(1)
C SET DATA SAMPLES INTO A REAL ARRAY
40 DO 42 MIN=N2,N3
  MFVE=MIN-N1
  A(MFVE)=IDAT(MIN)
42 CONTINUE
C SET NOISE SAMPLES INTO A REAL NUMBER ARRAY
  DO 130 J=1,N
  JEL=J+1
130 BMEAN(J)=IDAT(JEL)
  RETURN
C THE ID IS ERRONEOUS, IGNORE THE REST OF THE DATA
20 WRITE(6,130)ID,IDAT(1),IDAT(2)
130 FORMAT(40H ID WAS NOT CONSECUTIVE AND NOT=130050 ; ID=,3(I7,3X))
  KEOF=1
  RETURN
  END

```

```

CCCCCCCCCCCCCCCCCCCCCCCC---CALC2---CCCCCCCCCCCCCCCCCCCCCCCC
C   CALC2 IS A LIST OF CONSTANTS CALUATED FROM THE PROGRAMS
C   CALC AND ELDEN AND CONTAINS THE FUNCTION THAT CALCULATES THE
C   ELECTRON DENSITIES FOR THE PARTIAL-REFLECTION PROCESSING PROGRAMS.
C   THE PROGRAM CALC WRITES THIS PROGRAM.
CCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCC
SUBROUTINE CALC2(ARRAY,LL,LH,FD,IA)
DIMENSION ARRAY(21),RATIO2(21),FD(21)
C   GET THE PREDETERMINED CONSTANTS FOR THE RIGHT SEASON
IF(IA)200,300,100
C   CONSTANTS FOR THE SUMMER
100  RATIO2( 1)= 1.0731152
      RATIO2( 2)= 1.0778633
      RATIO2( 3)= 1.0841143
      RATIO2( 4)= 1.0909986
      RATIO2( 5)= 1.0990518
      RATIO2( 6)= 1.09901243
      RATIO2( 7)= 1.0880302
      RATIO2( 8)= 1.0864129
      RATIO2( 9)= 1.0768397
      RATIO2(10)= 1.0590323
      RATIO2(11)= 1.0495014
      RATIO2(12)= 1.0328141
      RATIO2(13)= 1.0262468
      RATIO2(14)= 1.0166022
      RATIO2(15)= 1.0117923
      RATIO2(16)= 1.0076428
      RATIO2(17)= 1.0044388
      RATIO2(18)= 1.0029128
      RATIO2(19)= 1.0015084
      RATIO2(20)= 1.0008443
      FD( 1)= 0.170327E-03
      FD( 2)= 0.243443E-03
      FD( 3)= 0.329813E-03
      FD( 4)= 0.427651E-03
      FD( 5)= 0.532098E-03
      FD( 6)= 0.627513E-03
      FD( 7)= 0.699730E-03
      FD( 8)= 0.744879E-03
      FD( 9)= 0.753246E-03
      FD(10)= 0.722645E-03
      FD(11)= 0.660879E-03
      FD(12)= 0.579908E-03
      FD(13)= 0.492934E-03
      FD(14)= 0.405772E-03
      FD(15)= 0.328260E-03
      FD(16)= 0.258602E-03
      FD(17)= 0.201748E-03
      FD(18)= 0.155429E-03
      FD(19)= 0.118415E-03
      FD(20)= 0.915812E-04
      GO TO 400
C   CONSTANTS FOR THE WINTER
200  RATIO2( 1)= 1.0682487
      RATIO2( 2)= 1.0890572
      RATIO2( 3)= 1.0699695
      RATIO2( 4)= 1.0979884
      RATIO2( 5)= 1.0692204
      RATIO2( 6)= 1.0854059
      RATIO2( 7)= 1.0770541
      RATIO2( 8)= 1.0745673
      RATIO2( 9)= 1.0665843
      RATIO2(10)= 1.0531266
      RATIO2(11)= 1.0458097
      RATIO2(12)= 1.0329590
      RATIO2(13)= 1.0286678
      RATIO2(14)= 1.0167225
      RATIO2(15)= 1.0127123
      RATIO2(16)= 1.0076428
      RATIO2(17)= 1.0046818
      RATIO2(18)= 1.0023713
      RATIO2(19)= 1.0019426
      RATIO2(20)= 1.0007090
      FD( 1)= 0.237897E-03
      FD( 2)= 0.319960E-03
      FD( 3)= 0.410827E-03
      FD( 4)= 0.502861E-03
      FD( 5)= 0.592594E-03
      FD( 6)= 0.662825E-03
      FD( 7)= 0.720342E-03
      FD( 8)= 0.750909E-03
      FD( 9)= 0.751034E-03
      FD(10)= 0.720293E-03
      FD(11)= 0.663644E-03

```



```

C*****-PROC73-*****
C      PROC73 EVALUATES COLLECTED PARTIAL-REFLECTION DATA AND
C PRINTS OUT THE ELECTRON DENSITY.  PROC73 USES THE FOLLOWING PROGRAMS:
C      HEAD---SETS UP AND PRINTS THE HEADING (FORTRAN)
C      DINIT---INITIALIZES THE STORAGE DEVIVE (MACRO)
C      FSTAT---LOCATES THE DATA FILE (MACRO)
C      SEEK---FINDS THE FILE ON THE STORAGE DEVICE (MACRO)
C      DRD73--SETSSAMPLES INTO THE REAL # ARRAY (FORTRAN)
C      CALC2---CALCULATES THE ELECTRON DENSITY (FORTRAN)
C*****
C
C      INTEGER DATIN
C      DIMENSION FNAM(2),AO(21),AX(21),AVAO(21),AVAX(21),
C      IXO(21),IRJ(21),BNO(5),BNX(5),EL(21)
C      WRITE(6,105)
105  FORMAT(48H TYPE IN SEASON--(1) FOR SUMMER, (-1) FOR WINTER
C      115H, (0) OTHERWISE)
C      READ(4,200)IA
200  FORMAT(I2)
C      DATIN=2
10  CALL HEAD(0)
C INITIALIZE VARIABLES
12  SNO=0.
C      SNX=0.
C      IR=0
C      IRN=0
C      BMO=0.
C      BMX=0.
C      DO 16 I=1,21
C      AVAO(I)=0.
C      AVAX(I)=0.
16  IRJ(I)=0
C      DO 17 I=1,4
C      BNO(I)=0.
C      BNX(I)=0.
17  INITIALIZE TAPE STORING THE DATA
C      CALL DINIT
C      GET THE DATA FILE NAME
C      WRITE(6,20)
20  FORMAT(15H WHICH DATAFILE)
C      READ(4,30)FNAM
30  FORMAT(2A5)
C      CHECK THE VALIDITY OF THE NAME GIVEN
13  CALL FSTAT(DATIN, FNAM, LOG)
C      IF (LOG.NE.0)GO TO 43
C      WRITE(6,35)FNAM
35  FORMAT(6H FILE ,2A5,19H NOT FOUND ON DAT 2)
C      30 TO 10
C      FIND LOCATION OF FILE ON THE TAPE
40  CALL SEEK(DATIN, FNAM)
C      GET THE MAXIMUM ALLOWABLE NOISE
C      WRITE(6,57)
57  FORMAT(14H MAXIMUM NOISE)
C      READ(4,56)BMXNS
56  FORMAT(F10.0)
C      IF (BMXNS.GE.510.)BMXNS=400.
C      FOR SPEED USE THE SQUARE OF THE MAX. ALLOW. NOISE TIMES 5
C      DUM4=BMXNS*5.
C      BMXNS=BMXNS/DUM4
19  KE0FO=0
C      KE0FX=0
C      ID=0
C      GET ONE SET OF 26 ORDINARY SAMPLES
48  CALL DRD73(AO,BNO, IERR, ID,KE0FO)
C      IF (KE0FO.EQ.1)GO TO 50
C      GET ONE SET OF 26 EXTRAORDINARY SAMPLES
C      CALL DRD73(AX,BNX, IERR, ID,KE0FX)
C      IF (KE0FX.EQ.1)GO TO 49
C      GET THE SUM SQUARED OF THE NOISE
C      BMEANO=0.
C      BMEANX=0.
C      DO 440 I=1,5
C      BMEANO=BMEANO+BNO(I)*BNO(I)
440  BMEANX=BMEANX+BNX(I)*BNX(I)
C      CHECK FOR SETS OF SAMPLES THAT ARE TOO NOISY
C      IF (BMEANO.GT.BMXNS.OR.BMEANX.GT.BMXNS)GO TO 510
C      SUM THE SQUARED NOISE SAMPLES FOR THE
C      LAST FOUR NOISE SAMPLES PER 25 TOTAL SAMPLES
C      BMO=BMO+BMEANO-BNO(I)*BNO(I)
C      BMX=BMX+BMEANX-BNX(I)*BNX(I)
730  DO 47 I=1,21
C      IF(AO(I).GE.510..OR.AX(I).GE.510.)GO TO 46

```



```

.TITLE READ DATA IN DUMP MODE
/XXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXX
/ .READ, DUMP MODE FROM DECTAPE ON A VARIABLE .DAT SLOT
/FILLS 252 DEC WORD BUFFER AND OUTPUTS 26
/WORDS TO ARRAY IDAT EVERY TIME CALLED.
/THESE ARE UNPACKED FROM 18 WORDS OF THE BUFFER.
/ IDAT: WORD 1          I.D. #
/          WORD 2-6     NOISE SAMPLES
/          WORD 7-27    DATA
/ NEGFI: SET IF A NEGATIVE NUMBER WAS IN THE DATA
/
.GLOBL DINIT,DUMPT,.DA
DUMP=4          /TYPE OF I/O MODE
DATIN=2        / .DAT SLOT TO READ DATA FROM
NSAMP=32       /# OF SAMPLES PER SET
NSAMP=NSAMP/2+1 /SIZE OF ONE SET PACKED DOUBLE
DTBLK=374      /SIZE OF ONE BLOCK OF STORAGE
DATSTR=DTBLK/NSAMP /# OF SETS PER ONE BLOCK OF STORAGE
NDPC=NSAMP/2   /NUMBER OF STORED DATA PAIRS PER SET

.DIODEV DATIN
DINIT 0
.INIT DATIN,0,DINIT /INITIALIZE DEVICE STORING THE DATA
LAW -1 /PREPARE TO READ
DAC CTT9 / IN ONE BLOCK OF DATA
LAC BUF31 /RESET THE BUFFER POINTER WITH
DAC POINT2 / THE ADDR. OF BUF3
LAC (ISZ SWITC /PREPARE TO READ TWO
DAC LCA / DUMMY BLOCKS
JMP* DINIT /END OF INITIALIZATION
DUMPT 0
JMS* .DA /PICKUP ADDR OF ADDR
JMP .+3 /OF ARRAY
A2 0
FLAG 0 /SET ON NEG #
LAC* A2 /GET ADDR.
DAC A2 / OF ARRAY
LAW -NDPC /SET COUNTER OF DATA TO BE
DAC COUNT / PROCESSED
ISZ CTT9 /GET POINTER
JMP LBB /NO, CONTINUE WITH PRESENT SET OF DATA
LAW -DATSTR /RESET COUNTER TO THE NUMBER
DAC CTT9 / OF SETS PER BLOCK OF STORAGE
LBA .READ DATIN,DUMP,BUF3,DTBLK /GET 1 BLOCK OF DATA
C14 .WAIT DATIN
LCA ISZ SWITC /INITIALLY READ TWO DUMMY BLOCKS
JMP LBA
LAW -3 /RESET CONTROL TO READ
DAC SWITC / TWO DUMMY BLOCKS
LAC (JMP LCB /READ ONE BLOCK OF DATA
DAC LCA / AT A TIME
LCB LAC BUF31 /GET ADDRESS OF BUF3
DAC POINT2 / POINT2 TO BUF3
LBB LAC* POINT2 /GET THE ID (FIRST WORD IN DATA SET)
SAD SEVN /END OF FILE ID?
JMP ENF /YES, RESET PARAM'S AND CLOSE FILE
LAC* POINT2 /GET ID AND PUT
DAC* A2 / INTO THE FORTRAN ARRAY
ISZ A2 /GO TO NEXT ADDR. OF THE ARRAY
LOOP ISZ POINT2 /GO TO NEXT DATA WORD
LAW -2 /PREPARE TO UNPAC
DAC TC10 / TWO DATA WORDS
LAC* POINT2 /GET DATA WORDS FROM BUF3
SWHA /FIRST WORD IN LEFT HALF
UNPLP AND (777 /SAVE ONE DATA WORD
SNA /CHECK FOR NEG. NUMBER
ISZ* FLAG /SET IF NEG. NUMBER FOUND
DAC* A2 /LOAD # INTO FORTRAN ARRAY
LAC* POINT2 /GET DATA WORD AGAIN
ISZ A2 /GO TO NEXT LOC. IN ARRAY
ISZ TC10 /UNPACED TWO WORDS?
JMP UNPLP /NO, LOOP AROUND
ISZ COUNT /YES, HAS 34 DATA WORDS BEEN UNPACKED?
JMP LOOP /NO, REPEAT UNPACKING PROCESS
OUT2 ISZ POINT2 /YES, GO TO NEXT ID
JMP* DUMPT /RETURN
SWITC LAW -3
COUNT LAW -NDPC
CTT9 LAW -DATSTR
SEVN 376002
TC10 0
POINT2 .DSA BUF3
BUF31 .DSA BUF3
BUF3 .BLOCK DTBLK
/-----

```

```

/ END OF FILE ROUTINE
ENF   LAC   SEVN   /SET LAST ID TO
      DAC*  A2     / 130050 DECIMAL
C15   .CLOSE DATIN /CLOSE FILE
RT2   JMP*  DUMPT  /RETURN TO PROC. PROGRAM
/
/XXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXXX
      .END

```

```

C***** CALC *****
C FROM GIVEN COLLISION FREQUENCIES, CALC ALONG WITH ELDEN
C CALCULATES THE CONSTANT VALUES USED IN THE ELECTRON DENSITY
C EQUATION GIVEN BY PIRNAT IN AERONOMY REPORT 29 AND WRITES THE
C PROGRAM CALC2 WHICH CALCULATES THE ELECTRO DENSITIES FOR THE
C PARTIAL-REFLECTION PROGRAMS.
C*****
      DIMENSION ARRAY(21),P(21),R(3),CF(3),EL(20),CALC2(2)
      DATA CALC2(1),CALC2(2)/5HCALC2,4H SRC/
      IDAT=1
      LABL=0
C COLLISION FREQUENCY PROFILES
C COLLISION FREQUENCY PROFILE FOR THE SUMMER
100 P(1)=192.3
      P(2)=156.9
      P(3)=127.5
      P(4)=102.7
      P(5)=82.37
      P(6)=66.25
      P(7)=52.50
      P(8)=41.66
      P(9)=32.81
      P(10)=25.84
      P(11)=20.1
      P(12)=15.53
      P(13)=11.89
      P(14)=9.057
      P(15)=6.817
      P(16)=5.399
      P(17)=3.827
      P(18)=2.862
      P(19)=2.124
      P(20)=1.563
      P(21)=1.180
C WRITE THE PROGRAM HEADING ONTO TAPE
      CALL ENTER(IDAT,CALC2)
      WRITE(IDAT,10)
10  FORMAT(69H CCCCCCCCCCCCCCCCCCCCCCCCCCCCC---CALC2---CCCCCCCCC
1CCCCCCCCCCCCCCCCCCCC/59H C CALC2 IS A LIST OF CONSTANTS
2 CALUATED FROM THE PROGRAMS/63H C CALC AND ELDEN AND COVT
3AINS THE FUNCTION THAT CALCULATES THE/69H C ELECTRON DENSITIES
4 FOR THE PARTIAL-REFLECTION PROCESSING PROGRAMS./40H C THE PRO
5GRAM CALC WRITES THIS PROGRAM./69H CCCCCCCCCCCCCCCCCCCCCC
6CCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCCC/37H SUBROUTINE
7 CALC2(ARRAY,LL,LH,FD,IA)/39H DIMENSION ARRAY(21),RATIO2
8(21),FD(21)/57H C GET THE PREDETERMINED CONSTANTS FOR THE
9 RIGHT SEASON/19H IF(IA)200,300,100/
128H C CONSTANTS FOR THE SUMMER)
      GO TO 400
C COLLISION FREQUENCY PROFILE FOR THE WINTER
200 P(1)=133.5
      P(2)=107.8
      P(3)=87.12
      P(4)=70.04
      P(5)=56.30
      P(6)=45.28
      P(7)=36.55
      P(8)=29.32
      P(9)=23.52
      P(10)=18.80
      P(11)=14.97
      P(12)=11.99
      P(13)=9.561
      P(14)=7.541
      P(15)=6.008
      P(16)=4.748
      P(17)=3.758
      P(18)=2.941
      P(19)=2.321

```

```

P(20)=1.810
P(21)=1.431
WRITE(IDAT,11)
11  FORMAT(11H          GO TO 400/
128H C  CONSTANTS FOR THE WINTER)
GO TO 400
C  COLLISION FREQUENCY PROFILE FOR THE EQUINOX
300  P(1)=160.2
P(2)=130.2
P(3)=105.3
P(4)=84.90
P(5)=68.25
P(6)=54.75
P(7)=43.57
P(8)=34.31
P(9)=27.07
P(10)=21.32
P(11)=16.82
P(12)=13.26
P(13)=10.33
P(14)=8.062
P(15)=6.246
P(16)=4.835
P(17)=3.758
P(18)=2.915
P(19)=2.260
P(20)=1.733
P(21)=1.359
WRITE(IDAT,12)
12  FORMAT(11H          GO TO 400/
128H C  CONSTANTS FOR EQUINOX)
C  SET THE COLLISION FREQUENCIES TO THE RIGHT ORDER OF MAGNITUDE
400  DO 401 I=1,21
P(I)=P(I)*(10.**5)
401  CONTINUE
C  STATEMENT LABEL FOR THE NEW PROGRAM
LABL=LABL+100
JI=JI+4
K=0
DO 20 I=1,20
K=K+1
CF(I)=P(K)
CF(2)=P(K+1)
C  CALCULATE CONSTANTS FOR THE ELECTRON DENSITY EQUATION
20  CALL ELDEN(R,CF,ARRAY(1),ARRAY(I+1),EL(I))
ARRAY(I)=ARRAY(2)/ARRAY(1)
C  WRITE FIRST CONSTANT WITH A STATEMENT LABEL
WRITE(IDAT,405)LABL,ARRAY(1)
405  FORMAT(1H ,13,12H          RATIO2( 1)=,F10.7)
C  WRITE THE REST OF THE RO AND RX CONSTANTS
DO 25 I=2,20
ARRAY(I)=ARRAY(I+1)/ARRAY(I)
WRITE(IDAT,410) I,ARRAY(I)
410  FORMAT(9H          RATIO2(,12,2H)=,F10.7)
C  WRITE THE CONSTANT DENOMINATORS
DO 30 I=1,20
WRITE(IDAT,420) I,EL(I)
420  FORMAT(5H          FD(,12,2H)=,E13.5)
C  CALCULATE THE REST OF THE CONSTANTS
IF(JI.LT.5)GO TO 200
IF(JI.LT.10)GO TO 300
C  WRITE THE REST OF THE PROGRAM TO CALCULATE ELECTRON DENSITIES
WRITE(IDAT,40)
40  FORMAT(18H 400 DO 10 I=LL,LH/48H          IF(ARRAY(I).EQ.0..OR.
1ARRAY(I+1).EQ.0.)GO TO 50/59H C  THE FUNCTION FOR THE CALCULA
TION OF ELECTRON DENSITIES/
351H          FD(I)=ALOG((ARRAY(I)/ARRAY(I+1))*RATIO2(I))/FD(I)/
410H          GO TO 10/12H 50 FD(I)=0./12H 10 CONTINUE/8H          RETURN)
CALL CLOSE(IDAT)
STOP
END

```

```

C:-----SUBROUTINE ELDEN-----
C  DURING DATA PROCESSING THERE ARE ONLY 2 VARIABLES
C  FOR EACH HEIGHT (AO AND AX). THE EQUATION FOR THE
C  ELECTRON DENSITY AS GIVEN BY BIRLY (1971) IS:
C  ED=LN((AX(1)/AO(1))/(RX(1)/RO(1)))/((AX(2)/AO(2))/(RX(2)/RO(2)))/FD
C  WHERE LN IS THE NATURAL LOG AND 1 AND 2 ARE HEIGHT 1 AND 2
C  SUBROUTINE ELDEN CALCULATES THE CONSTANTS RX,RO,AND FD
C  FOR EACH HEIGHT.
C:-----

```

```

SUBROUTINE ELDEN(AXBYAO,GNU,RXRO1,RYRO2,FD)
DIMENSION AXBYAO(3),RXBYRO(3),RX(3),RO(3),GNU(3),RATIO(3)
C APPROX INTEGRAL PARAMETERS
A4=2.3983474E-2
A3=1.1287513E+1
A2=1.1394160E+2
A1=2.4653115E+1
B6=1.8064128E-2
B5=9.3877372
B4=1.4921254E+2
B3=2.8958085E+2
B2=1.2049512E+2
B1=2.4656819E+1
D3=1.1630641
D2=1.6901002E+1
D1=6.6945939
E5=4.3605732
E4=6.4093464E+1
E3=6.8920505E+1
E2=3.535257E+1
E1=6.6314497
AXBYAO(3)=0
C GNU(3) IS MEAN COLLISION FREQUENCY AT THE INTERMEDIATE HEIGHT
C CALCULATE C INTEGRALS AT BOTH HEIGHTS AND FOR AVERAGE GNU
DUM=GNU(1)+GNU(2)
GNU(3)=0.5*DUM
DO 22 K=1,3
O=(2.59614E+7)/GNU(K)
X=7.3886E+6/GNU(K)
CTN=O*(O*(O*(O+A1)+A2)+A3)+A4
CTD=O*(O*(O*(O*(O+B1)+B2)+B3)+B4)+B5)+B6
CTO=CTN/CTD
CTYN=X*(X*(X*(X+A1)+A2)+A3)+A4
CTXD=X*(X*(X*(X*(X+B1)+B2)+B3)+B4)+B5)+B6
CTX=CTYN/CTXD
CFO=(O*(O*(O+D1)+D2)+D3)/(O*(O*(O*(O+E1)+E2)+E3)+E4)+E5)
CFX=(X*(X*(X+D1)+D2)+D3)/(X*(X*(X*(X+E1)+E2)+E3)+E4)+E5)
C CALCULATE RATIOS
RX(K)=SQRT((X*CTX)**2+(2.5*CFX)**2)
RO(K)=SQRT((O*CTO)**2+(2.5*CFO)**2)
RXBYRO(K)=RX(K)/RO(K)
RATIO(K)=AXBYAO(K)/RXBYRO(K)
22 CONTINUE
C CALCULATE FD FROM FINAL VALUES OF DO LOOP
FO=(5.*3.1824E+3*CFO)/(4.*3.0E+8*GNU(3))
FX=(5.*3.1824E+3*CFX)/(4.*3.0E+8*GNU(3))
FD=(FX-FO)*3.0E+9
RXRO1=RXBYRO(1)
RYRO2=RXBYRO(2)
RETURN
END

```

```

C*****CHECK*****
C PROGRAM READS IN THE NUMBER OF SAMPLES ASKS FOR BY OPERATOR.
C IF THE NUMBER IS ZERO, 31 NUMBERS ARE READ IN AND SET UP AS PARTIAL
C REFLECTION DATA IS (I.E. 5 NOISE SAMPLES AND 21 DATA POINTS & 5 EXTRA)
C DATA IS PRINTED OUT IN THE FORM OF ONE NUMBER PER HEIGHT AFTER EACH
C GROUP OF 26 SAMPLES ARE READ IN. THIS HAS BEEN USE TO CHECK THE
C RECIEVER AND A/D CONVERTER AGAINST THE REYNOLDS SYSTEM AND TO SEE
C IF EVERYTHING IS OPERATING AS IT SHOULD. IF THE NUMBER READ IN IS
C NOT ZERO, THAT NUMBER OF SAMPLES ARE READ FROM THE A/D CONVERTER
C AND AN AVERAGE OF ALL THE NUMBERS ARE TAKEN AND PRINTED OUT.
C*****

```

```

C DIMENSION IA(50),RA1(50),RA2(50)
MAX=50
I31=31
WRITE (6,110)
110 FORMAT(10H ADC CHECK)
C DEFAULT VALUE FOR THE # OF SAMPLES = 31
4 NS=I31
C READ # OF SAMPLE TO BE READ FROM A/D CONVERTER
5 READ (4,210) IDV
210 FORMAT(I5)
IF(IDV.NE.0) NS=IDV
IF(NS.NE.I31)GO TO 50

```

```

C   FOR 31 NUMBERS READ IN, THE FORM USED IS 2 SETS OF 31 SAMPLES
C   AS IN THE PARTIAL REFLECTION COLLECTION
      I1=0
25  ICH=0
      I1=I1+1
      CALL INPAD(IA,NS,ICH)
6   IF(ICH.EQ.0) GO TO 6
      IF(I1.GT.1)GO TO 11
      DO 13 I=1,I31
C   CONVERSION ALGORITHM FOR A/D CONVERTER NEG. #'S TO COMPUTER
C   NEGATIVE NUMBERS
      IF(IA(I).GT.511)IA(I)=3072+(4096+32768)*7+IA(I)
13  RA1(I)=FLOAT(IA(I))/.511
      GO TO 25
11  WRITE(6,101)
C   DO LOOP FOR SECOND SET OF NUMBERS READ IN
      DO 15 I=1,I31
      IF(IA(I).GT.511)IA(I)=3072+(4096+32768)*7+IA(I)
      HT=45.+FLOAT(I-1)*1.5
      IF(I.EQ.5.OR.I.EQ.11)WRITE(6,105)
      IF(I.EQ.11)WRITE(6,106)
      RA2(I)=FLOAT(IA(I))/.511
C   WRITE OUT THE NUMBERS IN AN ORDERLY WAY
15  WRITE(6,100) HT,RA1(I),RA2(I)
100 FORMAT(3X,F4.1,4HKM ,F5.0,4HMV ,F5.0,2HMV)
101 FORMAT(6H NOISE)
105 FORMAT(25H -----)
106 FORMAT(5H DATA)
      GO TO 5
C   THE FOLLOWING DUMPS THE AVERAGE OF THE A/D CONVERTER NUMBERS
C   AND ALSO GIVES THE VALUE IN MILLIVOLTS
50  INS=(NS+MAX-1)/MAX
      TNS=INS*MAX
      DO 60 J=1,INS
      ICH=0
      CALL INPAD(IA,MAX,ICH)
20  IF(ICH.EQ.0) GO TO 20
      DO 55 I=1,MAX
      IF(IA(I).GT.511)IA(I)=3072+(4096+32768)*7+IA(I)
55  AV=AV+FLOAT(IA(I))
60  CONTINUE
      AV=AV/TNS
      AVV=AV/.512
120 WRITE(6,120)AV,AVV
      FORMAT(9H AVERAGE=,F7.3,12H VOLTAGE=,F8.3,3H MV)
      GO TO 5
      STOP
      END

```

```

/   .TITLE A/D CONVERTER SERVICE ROUTINES FOR BG.-FG.
/   RPKM15 V3A SERVICE ROUTINES FOR THE HP 5610A A TO D
/   CONVERTER. THESE ROUTINES PERMIT INPUT OF ANY SPECIFIED
/   NUMBER OF SAMPLES INTO A CORE BUFFER. INPUT MAY BE OVER-
/   LAPPED WITH PROGRAM EXECUTION, AND CONTROL MAY BE RELINQUISHED
/   TO LOWER PRIORITY PROGRAMS WHILE DATA TRANSFER TAKES PLACE.
/   MACRO-15 CALLING SEQUENCE:
/   JMS INPAD
/   NUMBER OF SAMPLES REQUIRED
/   BUFFER ADDRESS
/   COMPLETION FLAG ADDRESS
/   REAL-TIME SUBROUTINE ADDRESS, PRIORITY LEVEL IN BITS 0-2
/   (EXAMPLE: 500000+RTSUBA)
/   (RETURNS HERE IMMEDIATELY)
/   IF THE 4TH WORD AFTER THE JMS IS 0, NO REAL-TIME SUBROUTINE
/   WILL BE ACTIVATED. NOTE: THE PRIORITY CODE FOR MAINSTREAM IS 1
/   THE COMPLETION FLAG IS CLEARED BY THE CALL TO INPAD,
/   AND SET TO +1 FOR NORMAL COMPLETION OR -1001 IF A DATA
/   TIMING ERROR OCCURS.
/

```

```

ADWCR=26 /A-D WORD COUNT
ADCAR=ADWCR+1 /AND CURRENT ADDRESS REGISTERS
.SCOM=100 /MONITOR'S COMMUNICATION AREA
ADWI=703724 /A-D CONVERTER WRITE INITIALIZE
ADSO=703701 /SKIP ON WORD COUNT OVERFLOW
ADST=703721 /SKIP ON DATA TIMING ERROR
ADCO=703704 /CLEAR OVERFLOW FLAG
ADCT=703744 /CLEAR TIMING FLAG

```



```

/ ENTRY POINT FOR A-D INTERFACE INITIALIZATION
/
INPAD      .GLOBL  INPAD,.DA
INPAD      0
INPAD      JMS*   .DA
INPAD      JMP    .+4
INAR       0
INWC       0
INFLAG     0
INR        JMP    INSET          /REPLACED BY "LAC*   INWC"
INR        TCA
INR        DAC*   (ADWCR)        /SET WORD COUNT
INR        LAW    -1
INR        IAD*   INAR          /BUFFER ADDRESS -1
INR        DAC*   (ADCAR)        / TO CURRENT ADDRESS REG.
INR        DZM*   INFLAG        /CLEAR FLAG
INR        DZM*   INSUB#        /CLEAR REAL-TIME SUBROUTINE
INR        ADWI   /INITIALIZE INTERFACE
INR        JMP*   INPAD          /RETURN
/
/ THE FOLLOWING CODE IS EXECUTED ONLY ONCE
INSET      LAC*   (.SCOM+55)      /GET ENTRY POINT ADDRESS OF .SETUP
ADSV      DAC    :
ADSV      JMS*   :-1            /CALL .SETUP TO CONNECT ADINT TO API
ADSV      ADSO
ADSV      ADINT
ADSV      DZM*   (204
ADSV
ADSV      LAC    (LAC*   INWC
ADSV      DAC    INR            /MODIFY INSTRUCTION
ADSV      JMP    INR            / AND JUMP TO IT
/
/
/ INTERRUPT SERVICE ROUTINE. EXECUTED IMMEDIATELY AFTER COMPLETION
/ OF DATA TRANSFER. DETERMINES STATUS OF A-D INTERFACE, SETS
/ COMPLETION FLAG AND ACTIVATES REAL-TIME SUBROUTINE.
/ RUNS AT API LEVEL 0.
/
ADINT      0
ADINT      DBA          /PAGE ADDRESSING MODE
ADINT      DAC    ADSVA    /SAVE AC
ADINT      ADST        /TIMING ERROR?
ADINT      SKP!CLAIAC  /NO,+1 TO AC
ADINT      LAW    -1001    /YES, ERROR CODE
ADINT      DAC*   INFLAG  /SET FLAG
ADINT      ADCO        /CLEAR
ADINT      ADCT        / INTERFACE FLAGS
ADINT      LAC    ADSVA    /RESTORE AC
ADINT      DBR        /SET TO LEAVE HARDWARE API LEVEL
ADINT      JMP*   ADINT    /RETURN TO INTERRUPTED PROGRAM
ADINT      .END

```