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T. Borkovits[1,2,3]⋆, S. A. Rappaport[4], P. F. L. Maxted[5] I. Terentev[6], M. Omohundro[6], R. Gagliano[7], T. Jacobs[8], M. H. Kristiansen[9,10], D. LaCourse[11], H. M. Schwengeler[6], A. Vanderburg[12], M. G. Blackford[13]

[1] *Baja Astronomical Observatory of University of Szeged, H-6500 Baja, Szegedi út, Kt. 766, Hungary*
[2] *Konkoly Observatory, Research Centre for Astronomy and Earth Sciences, H-1121 Budapest, Konkoly Thege Miklós út 15-17, Hungary*
[3] *ELTE Gothard Astrophysical Observatory, H-9700 Szombathely, Szent Imre h. u. 112, Hungary*
[4] *Department of Physics, Kavli Institute for Astrophysics and Space Research, M.I.T., Cambridge, MA 02139, USA*
[5] *Astrophysics Group, Keele University, Staffordshire, ST5 5BG, UK*
[6] *Citizen Scientist, c/o Zooniverse, Department of Physics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford, OX1 3RH, UK*
[7] *Amateur Astronomer, Glendale, AZ 85308*
[8] *Amateur Astronomer, 12812 SE 69th Place Bellevue, WA 98006, USA*
[9] *Brorfelde Observatory, Observator Gyldenkernes Vej 7, DK-4340 Tølløse, Denmark*
[10] *DTU Space, National Space Institute, Technical University of Denmark, Elektrovej 327, DK-2800 Lyngby, Denmark*
[11] *Amateur Astronomer, 7507 52nd Place NE Marysville, WA 98270, USA*
[12] *Department of Astronomy, The University of Wisconsin-Madison, 475 N. Charter St., Madison, WI 53706, USA*
[13] *Congarinni Observatory, Congarinni, NSW, Australia*

ABSTRACT

BG Ind is a well studied, bright, nearby binary consisting of a pair of F stars in a 1.46-day orbit. We have discovered in the *TESS* lightcurve for TIC 229804573 (aka BG Ind) a second eclipsing binary in the system with a 0.53-day. Our subsequent analyses of the recent *TESS* and archival ground-based photometric and radial velocity data, reveal that the two binaries are gravitationally bound in a 721-day period, moderately eccentric orbit. We present the results of a joint spectro-photodynamical analysis of the eclipse timing variation curves of both binaries based on *TESS* and ground-based archival data, the *TESS* lightcurve, archival radial velocity data and the spectral energy distribution, coupled with the use of PARSEC stellar isochrones. We confirm prior studies of BG Ind which found that the brighter binary A consists of slightly evolved F-type stars with refined masses of 1.32 and 1.43 $M_\odot$, and radii of 1.59 and 2.34 $R_\odot$. The previously unknown binary B has two less massive stars of 0.69 and 0.64 $M_\odot$ and radii of 0.64 and 0.61 $R_\odot$. Based on a number of different arguments which we discuss, we conclude that the three orbital planes are likely aligned to within 17°.

Key words: binaries:eclipsing – binaries:close – stars:individual:BG Ind

1 INTRODUCTION

BG Ind (κ_1 Ind; HD 208496; TIC 229804573) is a bright, sixth magnitude eclipsing binary (EB) formed by two F-type stars. Its variability was reported first by Strohmeier et al. (1964), and its EB nature was found by Manfroid & Mathys (1984). At the same time, Andersen et al. (1984) obtained the first two spectrograms, and concluded that BG Ind is also a double-lined spectroscopic binary and calculated stellar masses and radii for the first time. The

first photometric lightcurve analysis was carried out by van Hamme & Manfroid (1988).

In the forthcoming decades several new photometric and spectroscopic observations were carried out. They are nicely summarized in Rozyczka et al. (2011), and therefore we do not repeat them here.

The most recent thorough spectroscopic and photometric analysis was carried out by Rozyczka et al. (2011). These authors analysed all the available lightcurves and radial velocity (RV) data including their own measurements. They performed extensive spectroscopic analyses to obtain accurate stellar temperatures, system abundances and then age and evolutionary status. We will compare their results with

⋆ E-mail: borko@electra.bajaobs.hu

our findings later in Sect. 4, and therefore, here we highlight only a few noteworthy details. First, they found that the more massive and larger star has the lower temperature[1], thereby indicating clearly that this component has already evolved away from the main sequence and is moving toward the subgiant regime. Second, they made attempts to resolve some problems with both the photometric phasing (already first noted in van Hamme & Manfroid 1988) and discrepancies in the systemic γ velocities obtained in the solutions of the RV curves measured during three highly different epochs by Andersen et al. (1984); Bakış et al. (2010); Rozyczka et al. (2011). However, they were not able arrive at any definitive conclusions regarding these inconsistencies.

We further note that BG Ind was included in the catalog of those detached eclipsing binaries for which the constituent masses and radii are known to at least 2% precision (Southworth 2015). And BG Ind was also selected for inclusion in the sample of 156 detached eclipsing binaries, which can be used as benchmarks for trigonometric parallaxes in the Gaia era (Stassun & Torres 2016). Finally, turning to the Gaia era, with the use of Gaia DR2 (Gaia collaboration 2018) and HIPPARCOS (van Leeuwen 2007) data, a significant proper motion anomaly was found that might indicate the presence of further, gravitationally bound components in the system (Brandt 2018; Kervella et al. 2019). At this point it should also be noted, that there is a remarkable discrepancy between the revised HIPPARCOS and Gaia EDR3 (Gaia collaboration 2020) parallaxes of BG Ind ($\pi_{\mathrm{HIP}} = 14.90 \pm 0.59$ mas vs $\pi_{\mathrm{EDR3}} = 19.44 \pm 0.52$ mas), which might be a further indicator of additional multiplicity in the system.

In this paper we confirm the—at least—quadruple nature of BG Ind. Using the high-precision *TESS* photometry with 2-min cadence we have discovered an obvious second eclipsing EB in the lightcurve of BG Ind with a period of 0.53 days[2]. Our comprehensive investigation of the *TESS* photometry, archival ground-based photometry and radial velocity curves, as well as the eclipse timing variations (ETV) data demonstrates that the two eclipsing binaries form a close 2+2 quadruple stellar system with a remarkably short outer period of ~ 2 years.

In Sect. 2 we describe all the available observational data and their preparation for the complex, joint photodynamical analysis which is discussed in Sect. 3. Then, the results are discussed and finally, summarized in Sects. 4 and 5.

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[1] In most binaries with unevolved stars the primary star is the more massive and the hotter star. In the A binary of this system (the dominant binary), the more massive star turns out to be the cooler of the two due to its evolution. We continue to refer to the more massive star as the 'primary' even though it is cooler. However, we will still refer to the deeper eclipse (i.e., when the primary eclipses the less massive (but hotter star) as the 'primary eclipse'. The usual naming convention still holds for the fainter binary B.

[2] This second eclipsing binary was also found independently by Eisner et al. 2021.

Table 1. Main properties of BG Ind from different catalogs.

Parameter	Value	References
RA	329.62537	1
DEC	-59.01201	1
μ_{RA} [mas yr^{-1}]	4.96 ± 0.35	1
μ_{DEC} [mas yr^{-1}]	30.21 ± 0.53	1
π_{EDR3} [mas]	19.44 ± 0.52	1
π_{HIP} [mas]	14.90 ± 0.59	2
G	6.024606 ± 0.001610	1
G_{BP}	6.266601 ± 0.005111	1
G_{RP}	5.645544 ± 0.010610	1
T	5.6502 ± 0.0067	3
B	6.605 ± 0.022	3
V	6.130 ± 0.030	3
B_T	6.697 ± 0.014	4
V_T	6.195 ± 0.009	4
J	5.206 ± 0.020	5
H	4.993 ± 0.026	5
K	4.877 ± 0.026	5
W1	4.907 ± 0.215	6
W2	4.615 ± 0.092	6
W3	4.897 ± 0.014	6
W4	4.821 ± 0.028	6
FUV	15.216 ± 0.015	7
NUV	11.598 ± 0.002	7
$[M/H]$ [dex]	-0.30	8
Distance [pc]	51.0 ± 0.5	9

References. (1) Gaia EDR3 (Gaia collaboration 2020); (2) HIPPARCOS (revised) (van Leeuwen 2007); (3) TIC-8 catalog (Stassun et al. 2018); (4) Tycho-2 catalog (Høg et al. 2000); (5) 2MASS All-Sky Catalog of Point Sources (Skrutskie et al. 2006); (6) AllWISE catalog (Cutri et al. 2013); (7) GALEX-DR5 (GR5) (Bianchi et al. 2011) ;(8) Holmberg et al. (2009); (9) Bailer-Jones et al. (2018)

2 OBSERVATIONAL DATA

2.1 Catalog data

In Table 1, in addition to other catalog data, we collected the photometric passband magnitudes of the system from different surveys, e.g. Tycho-2 (Høg et al. 2000), 2MASS (Skrutskie et al. 2006), AllWISE (Cutri et al. 2013), GALEX (Bianchi et al. 2011) and Gaia (Gaia collaboration 2020). These will be used to construct the spectral energy distribution ('SED') of the system. In turn, the SED along with theoretical isochrones and the photodynamical model of the system provide an opportunity to determine the masses of the components in an astrophysical model-dependent way (see Sect. 3 for details). Together with the passband magnitudes given in Table 1, we list their uncertainties as tabulated in the given catalogs. For the SED analysis, however, we used a minimum uncertainty of 0.03 mag to avoid the overdominance of the extremely precise Gaia magnitudes and also to counterbalance the uncertainties inherent in our interpolataion method during the calculations of theoretical passband magnitudes that are part of the fitting process. Furthermore, similar to the approach followed by Stassun & Torres (2016) we omitted the GALEX near-UV magnitude from our analysis as a distinct outlier. K. Stassun (private communication) kindly called our attention to the fact that even the largest available NUV aperture is missing flux.
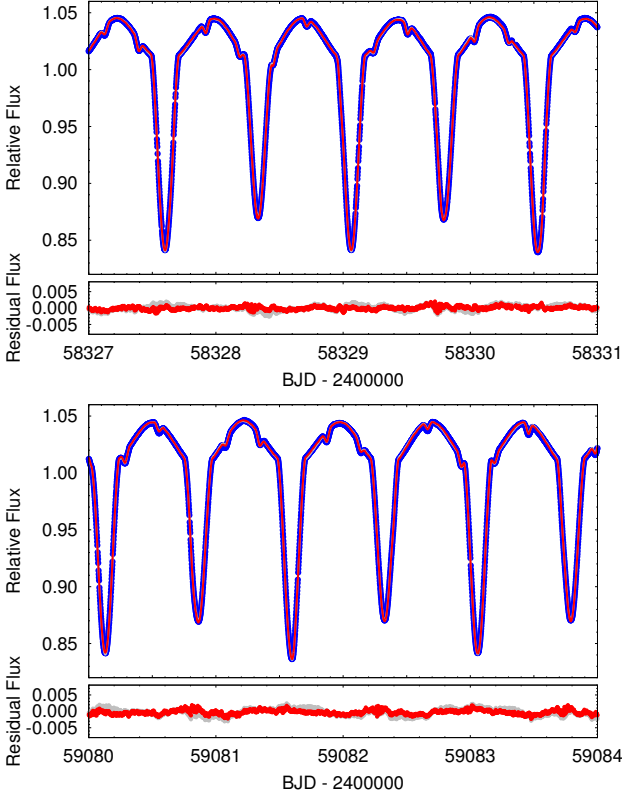
Figure 1. Two four-day sections from the beginning of Sector 1 and the end of Sector 28 SAP lightcurves of BG Ind (blue circles). The red and grey curves are spectro-photodynamical model solutions (see later, in Sect. 3). In the case of the red solution the small extra fluctuations of the lightcurve are probably due to the chromospheric/photospheric activities of the stars and were modelled mathematically with Fourier-harmonics simultaneously with the two-binary model, while the grey curve represent the pure two-binary part of the same solution. The residuals to the models are also shown below the lightcurves.

2.2 *TESS* photometry

The *TESS* space telescope (Ricker et al. 2015) has observed this target in 2-min cadence mode during Sectors 1, 27 and 28. We downloaded both the Simple Aperture Photometry (SAP) and the Pre-search Data Conditioning SAP (PDC-SAP) lightcurves from the MAST portal[3]. We used the SAP lightcurves for our study. Because the presence of the small extra dips belonging to the eclipses of the previously unknown binary B (see Fig. 1) was discovered shortly after the release of the data of the first four *TESS* sectors, our analyses were carried out mostly with the use of Sector 1 data. We also did use Sector 27 and 28 data, but mainly for the purpose of extending the interval of the ETV study. Since, in the case of the faint binary B, the only sources of ETV data are the three sectors of high-quality *TESS* data, the inclusion of these new observations into our analysis significantly improved the accuracy of the outer orbit solution (including the dynamically inferred mass of binary B).
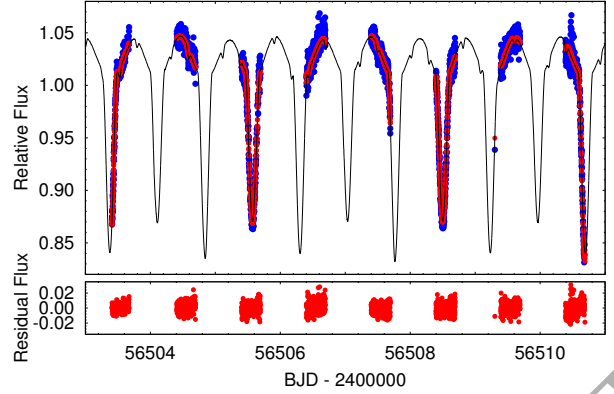
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[3] https://mast.stsci.edu

Figure 2. An eight-day-long section of the WASP observations of BG Ind (blue circles). Red dots represent the best-fit spectro-photodynamical model solution (see later, in Sect. 3) projected back to the epoch of each individual WASP measurement, while the black line shows the evenly phased model solution. The residuals curve is also shown below.

2.3 WASP photometry

BG Ind is one of millions of stars that have been observed as part of the WASP survey. The survey is described in Pollacco et al. (2006) and Collier Cameron et al (2006). From July 2012 the WASP-South instrument was operated using 85-mm, f/1.2 lenses and an r′ filter. With these lenses the image scale is 33 arcsec/pixel. Observations of BG Ind were obtained simultaneously in two cameras on WASP-South over three observing seasons, from 2012 July 3 to 2014 December 6. Fluxes are measured in an aperture with a radius of 132 arcsec for the 85-mm data and instrumental trends are removed using the SYSRem algorithm (Tamuz et al. 2005). Data points more than 5 standard deviations from a phase-binned version of the light curve were rejected and the entire night of data was rejected if more than 1/4 of the observations were identified as outliers based on this criterion.

An eight-day section of the WASP measurements is shown in Fig. 2. Note, we converted the original HJD(UTC) times of the WASP observations to BJD(TDT) for the forthcoming analyses. We also applied the same transformations for all the archival data that we describe below.

2.4 Other ground-based archive photometric data used for our analysis

We downloaded publicly available Strömgren u, v, b, y photometric observations from the ESO archive (Manfroid et al. 1991; Sterken et al. 1993). These observations were carried out between JDs 2 446 581 and 2 447 069. The lightcurves in each bandpass contain 175 measurements. We used these data primarily to determine additional times of eclipses in binary A. Unfortunately, however, the majority of the nightly observations contain only a few measurements, and we were therefore able to determine the mid-times of only two primary eclipses (see below, in Sect. 2.6) from this data set.

BG Ind was also observed by Jens Viggo Clausen and collaborators as part of their long-running observing pro-

gramme to measure absolute dimensions for solar-type stars in eclipsing binaries, carried out since 1994 at the Strömgren Automatic Telescope at ESO, La Silla. Unpublished data and manuscripts that were in preparation from this observing programme have been made available to one of us (PM), from which we extracted 5 more eclipse times (see again, in Sect. 2.6). The observing procedures and data reduction for these observations are similar to those described in Clausen et al. (2001).

2.5   Disentanglement of the lightcurves

For the combined analysis of all the observational data we used the original *TESS* timeseries, i.e., the net lightcurve of the two binaries together. However, at the start of the analysis, we found it worthwhile to disentangle the lightcurves of the two binaries, so we could examine each one separately. We have described this process in substantial detail in Powell et al. (2021). Therefore we review only the highlights here. First we folded and binned the Sector 1 (i.e. Year 1) *TESS* SAP lightcurve with the period of binary A into 1000 equal phase cells. However, while producing the fold for binary A, we excluded those data points which were recorded during the eclipses of binary B. Then the mean flux of each cell was rendered to the mid-phase value of that cell. In such a way we obtained a folded, binned, and averaged lightcurve of binary A (see upper panel of Fig. 3). Then this lightcurve was removed from the original sector 1 SAP lightcurve point by point in such a manner that the flux to be removed at the actual phase of any given data point was calculated with a three-point local Lagrange-interpolation from the folded, binned, and averaged lightcurve. As the result of this removal, we have obtained a new, residual time series which now mainly contains the light variations of binary B,[4] without the eclipses and ellipsoidal variations of binary A. Therefore, this lightcurve can be used for determining the mid-eclipse times of binary B.

In the next step we folded, binned, and averaged this residual lightcurve with the period of binary B (see middle panel of Fig. 3). Finally, we subtracted this folded, disentangled lightcurve of binary B from the original Sector 1 *TESS* SAP lightcurve, thereby obtaining a time series of binary A without the small distortions caused by binary B. We applied the same process to the Sector 27 and 28 (Year 3) SAP lightcurves, as well (see bottom panel of Fig. 3).

Regarding the WASP observations, we carried out a very similar process with the slight modification that, in this case, for the much smaller number of individual data points we applied binnings of 200 and 500 cells instead of 1000. We carried out the whole process separately for the three seasons of the WASP observations. Though the eclipsing signal of the faint binary B is not readily detected in the original WASP timeseries, we were able to see it clearly in our disentangled version (see Fig. 4).

Finally, in regard to disentangling the lightcurves, we

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[4] Note, for practical reasons, we added a constant flux to these time series in such a way that the flux of the very first data point retained the same value as in the original time series. In this manner, we replaced the varying light of the extracted binary with a constant extra light.
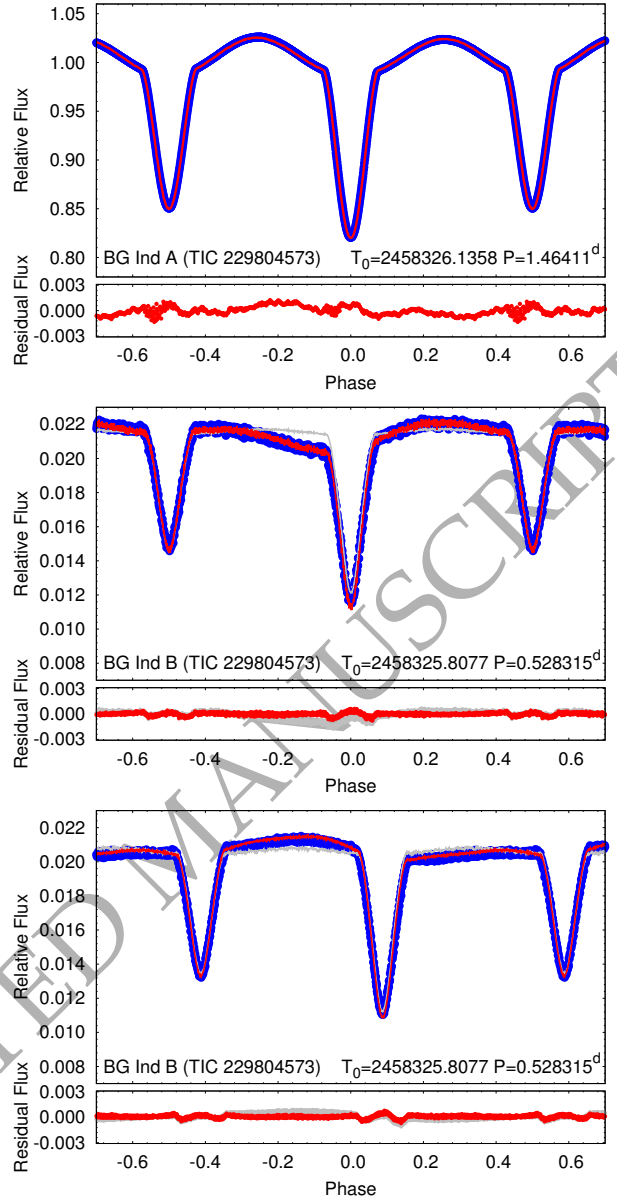
Figure 3. Folded, binned, and averaged *TESS* lightcurves of the two binaries of the quadruple system BG Ind. *Upper panel:* Sector 1 lightcurve of binary A (blue circles), together with the folded, binned, and averaged combined spectro-photodynamical model lightcurve (red curve, see later, in Sect. 3). *Middle and lower panels:* Year 1 (Sector 1) and Year 3 (Sector 27 and 28) lightcurves of binary B, respectively. As in the case of the binary A lightcurve in Fig. 1, the red solution curve exhibits some small extra fluctuations that are probably due to the chromospheric/photospheric activities of the stars (see Sect. 3 for details). These were modelled mathematically with Fourier-harmonics simultaneously with the two-binary model, while the thin grey curves represent the pure two-binary part of the same solution. The fold of the residuals to the models are also shown below the folded lightcurves.

Figure 4. Folded, binned, averaged WASP lightcurves of binary A and B for observing season 2012/2013. For illustrative purposes we phased both curves with the ephemeris calculated for Sector 1 *TESS* data. In such a manner, the shift of the primary and secondary eclipses from phases $0^{\rm p}0$ and $0^{\rm p}5$, respectively, i.e., the phasing problem mentioned in the Introduction, is clearly visible.

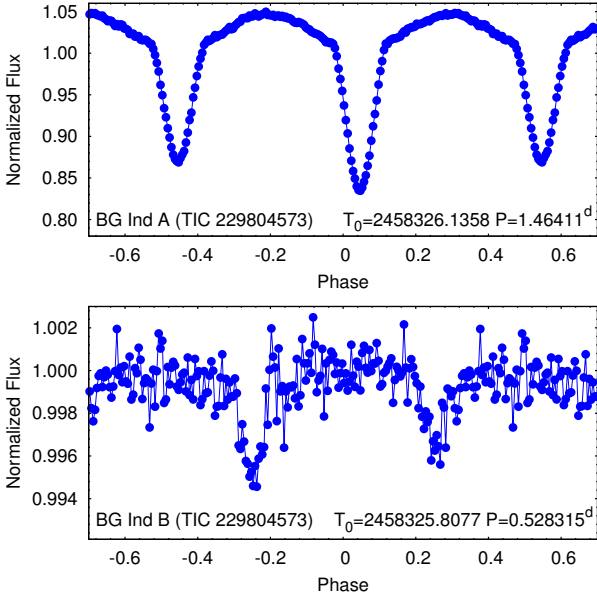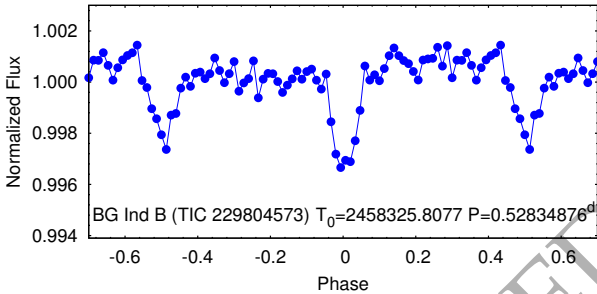**Figure 5.** The folded, binned, averaged WASP lightcurve of binary B obtained with the use of the second disentanglement method based on Fourier-filtering. The full, three-season WASP lightcurve was folded with the mean orbital period obtained from the ETV analysis (see Sect. 3).

have also used a second method which fits simultaneously for 50 harmonics of each of binaries A and B given their established periods. This technique, which is also described in detail in Powell et al. (2021), involves inverting a 201×201 matrix to solve for the linear coefficients to the 50 sines and 50 cosines for each of the two binaries. We find nearly perfect agreement for the disentangled *TESS* lightcurves from the two independent methods, and thus we do not show those results here. In the case of disentangling the WASP data, the results for binary B are actually somewhat improved using the Fourier approach and we show that lightcurve in Fig. 5 as well for comparison.

2.6 ETV data

2.6.1 TESS *ETV results*

In order to calculate accurate eclipse times from the *TESS* lightcurve we used the disentangled time series (see above in Sect. 2.5). The 91 eclipse times of binary A (including both primary and secondary eclipses) from Sectors 1, 27, and 28 are presented in Table 2. In Table 3 we list the eclipse times for binary B including a combined 259 primary and secondary eclipses.

2.6.2 *Ground-based ETV results*

We also utilized WASP and ESO (Manfroid et al. 1991; Sterken et al. 1993) data, including the unpublished observations of J. V. Clausen, as well, to calculate 85 additional eclipse times for binary A. Furthermore, a primary and a secondary eclipse of BG Ind were observed by one of us (MB) using a DSLR camera. Images were recorded in RAW format and the green, blue and red channels were extracted into separate images. The times of minimum were measured from each colour filter with the Peranso software[5] using a 5th order polynomial fit. The average of the mid-eclipse times were converted into BJD. Finally, we collected one other eclipse time from the the paper of van Hamme & Manfroid (1988) and converted it into BJD. All these eclipse times are tabulated in Table 4. The eclipses from binary B in the archival data were too weak to derive meaningful eclipse times.

2.6.3 *BG Ind ETV Results*

The overall ETV curves for BG Ind A and B are plotted in Fig. 6 along with the best-fit spectro-photodynamical model that is described in the next section. The ETV curve of BG Ind A exhibits a clearly cyclic pattern with a period of ~ 2 years. Even in the absence of any other indications of additional stars in the system, the most plausible explanation of this ETV behavior would be the light-travel time effect (LTTE) caused by a gravitationally bound, distant, third component. Therefore, we carried out, a preliminary, 'traditional' analysis of the ETV curves of binary A by fitting the LTTE-term with our analytic ETV-solver (Borkovits et al. 2015). We found that the very first eight ETV points deviate systematically from the LTTE solution. Therefore, we added a quadratic term to the analysis and obtained the following quadratic ephemeris:

$$
\begin{aligned}
T_{\rm pri} = &\ 2\,458\,326.13561(7) + 1.46406518(8) \times E \\
&+ 1.8(3) \times 10^{-10} \times E^2.
\end{aligned} \tag{1}
$$

We also tabulate the parameters of this preliminary LTTE solution in Table 5, and plot this simple model together with the spectro-photodynamical model, in Fig. 6.

Turning to the ETV points of binary B, they appear to be moving with the opposite phase to that of the bright binary A which makes it very likely that the two binaries form a bound, quite tight quadruple system. As we will discuss below in Sects. 3 and 4 our detailed analysis robustly confirms this hypothesis.

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[5] https://www.cbabelgium.com/peranso/

Table 2. Eclipse times of BG Ind binary A determined from *TESS* observations.

BJD −2 400 000	Cycle no.	std. dev. (d)	BJD −2 400 000	Cycle no.	std. dev. (d)	BJD −2 400 000	Cycle no.	std. dev. (d)
58325.403840	-0.5	0.000005	58349.561705	16.0	0.000005	59056.704428	499.0	0.000005
58326.135712	0.0	0.000004	58350.293558	16.5	0.000004	59057.436384	499.5	0.000005
58326.867912	0.5	0.000005	58351.025742	17.0	0.000006	59058.168508	500.0	0.000005
58327.599837	1.0	0.000003	58351.757721	17.5	0.000007	59058.900490	500.5	0.000006
58328.332109	1.5	0.000005	58352.489843	18.0	0.000004	59059.632635	501.0	0.000005
58329.063938	2.0	0.000004	59036.939185	485.5	0.000005	59060.364549	501.5	0.000007
58329.796203	2.5	0.000005	59037.670779	486.0	0.000005	59062.560877	503.0	0.000004
58330.528118	3.0	0.000004	59038.403292	486.5	0.000006	59063.292963	503.5	0.000005
58331.260286	3.5	0.000006	59039.134927	487.0	0.000004	59064.024900	504.0	0.000004
58331.992218	4.0	0.000004	59039.867263	487.5	0.000006	59064.757088	504.5	0.000004
58332.724303	4.5	0.000006	59040.598986	488.0	0.000005	59065.489098	505.0	0.000004
58333.456408	5.0	0.000004	59041.331371	488.5	0.000006	59066.221307	505.5	0.000005
58334.188441	5.5	0.000005	59042.063196	489.0	0.000005	59066.953158	506.0	0.000004
58334.920563	6.0	0.000005	59042.795556	489.5	0.000006	59067.685476	506.5	0.000006
58335.652617	6.5	0.000005	59043.527276	490.0	0.000005	59068.417320	507.0	0.000003
58336.384671	7.0	0.000004	59044.259642	490.5	0.000006	59069.149723	507.5	0.000007
58337.116773	7.5	0.000005	59044.991370	491.0	0.000005	59069.881452	508.0	0.000004
58337.848766	8.0	0.000005	59045.723835	491.5	0.000006	59070.613800	508.5	0.000007
58340.044827	9.5	0.000005	59046.455589	492.0	0.000005	59071.345545	509.0	0.000004
58340.777035	10.0	0.000004	59047.187905	492.5	0.000006	59075.737851	512.0	0.000004
58341.508927	10.5	0.000006	59047.919617	493.0	0.000004	59076.470079	512.5	0.000006
58342.241178	11.0	0.000004	59049.383671	494.0	0.000004	59077.201896	513.0	0.000004
58342.973078	11.5	0.000004	59050.116215	494.5	0.000006	59077.934281	513.5	0.000006
58343.705279	12.0	0.000004	59050.847785	495.0	0.000004	59078.666081	514.0	0.000005
58344.437009	12.5	0.000006	59051.580292	495.5	0.000006	59079.398485	514.5	0.000007
58345.169391	13.0	0.000005	59052.311934	496.0	0.000005	59080.130262	515.0	0.000005
58345.901131	13.5	0.000006	59053.044348	496.5	0.000007	59080.862593	515.5	0.000006
58346.633499	14.0	0.000005	59053.776090	497.0	0.000006	59081.594352	516.0	0.000004
58347.365392	14.5	0.000007	59054.508387	497.5	0.000006	59082.326656	516.5	0.000005
58348.097205	15.0	0.000074	59055.240325	498.0	0.000005	59083.058456	517.0	0.000004
58348.829518	15.5	0.000074	59055.972499	498.5	0.000004	59083.790713	517.5	0.000004

Notes. Integer and half-integer cycle numbers, as above, refer to primary and secondary eclipses, respectively.
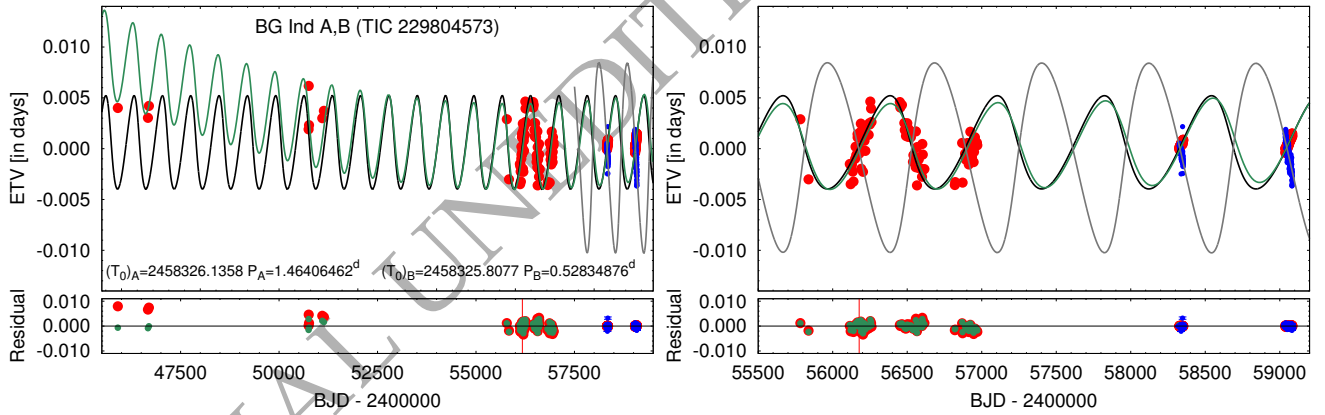
Figure 6. Eclipse timing variations of BG Ind. *Left panel* shows all the available observations, while in the *right panel* we zoom in on the regions of the better-covered WASP and *TESS* data. Larger red circles represent ETV points calculated from the observed eclipse events of binary A, while the smaller blue circles stand for the ETVs of binary B. Note, for simplicity, we do not separate primary and secondary eclipses. (The validity of this can easily be verified since both binaries have circular orbits and, furthermore, the primary and secondary eclipses within each binary can be calculated with the same accuracy due to their similar depths.) Black and grey lines stand for the combined spectro-photodynamical model ETV solution (Sect. 3) for binary A and B, respectively, while the green line denotes the preliminary, 'classic', analytic LTTE+quadratic ETV solution discussed in Sect. 2.6.3. (Note, for clarity, in the left panel, the ETV solution of binary B, i.e., the grey curve, is plotted only for the narrow interval around the *TESS* observations.) The residuals of the observed vs. modelled ETVs are plotted in the bottom panel. Here, as above, red and blue dots represent the residuals of binary A and B ETV points against the spectro-photodynamical model, while green dots stand for the residuals of binary A data against the analytic LTTE+quadratic ETV model.

Table 3. Eclipse times of BG Ind binary B determined from *TESS* observations.

BJD −2 400 000	Cycle no.	std. dev. (d)	BJD −2 400 000	Cycle no.	std. dev. (d)	BJD −2 400 000	Cycle no.	std. dev. (d)
58325.543752	-0.5	0.000062	58337.430905	22.0	0.000054	58350.902971	47.5	0.000085
58325.807997	0.0	0.000046	58337.693064	22.5	0.000104	58351.167488	48.0	0.000079
58326.072387	0.5	0.000072	58337.959793	23.0	0.000049	58351.431071	48.5	0.000088
58326.336575	1.0	0.000064	58338.222995	23.5	0.000055	58351.696086	49.0	0.000054
58326.600567	1.5	0.000072	58338.487016	24.0	0.000049	58351.959199	49.5	0.000097
58326.864689	2.0	0.000053	58339.808135	26.5	0.000070	58352.224055	50.0	0.000051
58327.128861	2.5	0.000066	58340.072263	27.0	0.000055	58352.488050	50.5	0.000058
58327.392666	3.0	0.000052	58340.336199	27.5	0.000079	58352.752014	51.0	0.000051
58327.657142	3.5	0.000077	58340.601135	28.0	0.000059	58353.016148	51.5	0.000071
58327.921372	4.0	0.000046	58340.864891	28.5	0.000102	59036.438310	1345.0	0.000073
58328.185258	4.5	0.000060	58341.129590	29.0	0.000056	59036.702049	1345.5	0.000075
58328.449821	5.0	0.000058	58341.392795	29.5	0.000076	59036.966367	1346.0	0.000071
58328.713709	5.5	0.000060	58341.657591	30.0	0.000068	59037.230717	1346.5	0.000068
58328.977792	6.0	0.000048	58341.921373	30.5	0.000092	59037.494934	1347.0	0.000061
58329.241446	6.5	0.000062	58342.185766	31.0	0.000078	59037.758430	1347.5	0.000073
58329.506583	7.0	0.000058	58342.450293	31.5	0.000057	59038.023264	1348.0	0.000068
58329.770573	7.5	0.000072	58342.714094	32.0	0.000053	59038.286595	1348.5	0.000054
58330.034386	8.0	0.000069	58342.978300	32.5	0.000077	59038.551986	1349.0	0.000089
58330.298768	8.5	0.000102	58343.242670	33.0	0.000051	59038.815125	1349.5	0.000088
58330.562823	9.0	0.000064	58343.506139	33.5	0.000082	59039.080123	1350.0	0.000068
58330.826710	9.5	0.000060	58343.770936	34.0	0.000070	59039.342745	1350.5	0.000095
58331.091287	10.0	0.000051	58344.034952	34.5	0.000081	59039.608385	1351.0	0.000062
58331.355655	10.5	0.000094	58344.299072	35.0	0.000080	59039.871119	1351.5	0.000082
58331.619925	11.0	0.000054	58344.562888	35.5	0.000099	59040.137118	1352.0	0.000074
58331.883907	11.5	0.000072	58344.827200	36.0	0.000046	59040.400668	1352.5	0.000074
58332.147667	12.0	0.000050	58345.091215	36.5	0.000073	59040.665367	1353.0	0.000086
58332.411725	12.5	0.000079	58345.355528	37.0	0.000049	59040.928474	1353.5	0.000086
58332.676452	13.0	0.000059	58345.619873	37.5	0.000074	59041.193535	1354.0	0.000075
58332.940235	13.5	0.000058	58345.884337	38.0	0.000053	59041.456972	1354.5	0.000074
58333.204648	14.0	0.000046	58346.148236	38.5	0.000069	59041.721550	1355.0	0.000070
58333.468688	14.5	0.000074	58346.412325	39.0	0.000057	59041.985219	1355.5	0.000102
58333.732522	15.0	0.000044	58346.676294	39.5	0.000100	59042.249607	1356.0	0.000060
58333.997359	15.5	0.000065	58346.940771	40.0	0.000071	59042.513262	1356.5	0.000073
58334.261381	16.0	0.000069	58347.204409	40.5	0.000097	59042.777673	1357.0	0.000064
58334.525599	16.5	0.000072	58347.468581	41.0	0.000263	59043.041442	1357.5	0.000112
58334.789683	17.0	0.000058	58347.736341	41.5	0.000426	59043.306328	1358.0	0.000066
58335.053444	17.5	0.000070	58347.997914	42.0	0.000197	59043.570070	1358.5	0.000101
58335.317858	18.0	0.000069	58348.526424	43.0	0.000392	59043.834600	1359.0	0.000077
58335.581721	18.5	0.000080	58349.053670	44.0	0.000225	59044.098701	1359.5	0.000065
58335.846166	19.0	0.000048	58349.316776	44.5	0.000139	59044.362603	1360.0	0.000078
58336.110542	19.5	0.000061	58349.582154	45.0	0.000055	59044.626253	1360.5	0.000079
58336.374765	20.0	0.000048	58349.845864	45.5	0.000060	59044.891275	1361.0	0.000081
58336.638703	20.5	0.000079	58350.110733	46.0	0.000050	59045.155782	1361.5	0.000080
58336.902547	21.0	0.000061	58350.374797	46.5	0.000088	59045.419161	1362.0	0.000068
58337.166804	21.5	0.000084	58350.639299	47.0	0.000067	59045.683175	1362.5	0.000077

*Notes.* Integer and half-integer cycle numbers, as above, refer to primary and secondary eclipses, respectively.

2.7 Radial Velocity data

We used three sets of radial velocity (RV) data for our analysis. These are as follows: (*i*) Bakış et al. (2010) have obtained 41 RVs between JDs 2 453 968 and 2 453 996 (i.e. in August 2006) with the High Efficiency and Resolution Canterbury University Large Echelle Spectrograph (HER-CULES) of the Department of Physics and Astronomy, New Zealand; (*ii*) an additional 23 RVs for both components of binary A between JDs 2 454 363 and 2 454 376 (September/October 2007) were collected by Rozyczka et al. (2011) with the fibre-fed Giraffe spectrograph on the 1.9-m Radcliffe telescope at the South African Astronomical Observatory; and (*iii*) finally, we found in the ESO publicly accessi-

ble archive[6] a number of spectra taken with the HARPS and FEROS spectrographs between JDs 2 453 191 and 2 456 910. From these data we determined an additional 54 RV points of both components of binary A.

To measure the RVs from these latter spectra we used the broadening function method (Rucinski 1992) as implemented in the software package RaveSpan (Pilecki et al. 2017). The template used for the analysis was a synthetic spectrum for a star with $T_{\text{eff}} = 7000$ K, $\log g = 3.5$. We used a simultaneous least-squares fit of two rotationally-broadened profiles to measure the radial velocities of the two

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[6] http://archive.eso.org/cms.html

Table 3. continued (*TESS* binary B eclipse times)

BJD −2 400 000	Cycle no.	std. dev. (d)	BJD −2 400 000	Cycle no.	std. dev. (d)	BJD −2 400 000	Cycle no.	std. dev. (d)
59045.946568	1363.0	0.000077	59058.099423	1386.0	0.000073	59070.513321	1409.5	0.000065
59046.210665	1363.5	0.000093	59058.362578	1386.5	0.000088	59070.778828	1410.0	0.000054
59046.476190	1364.0	0.000080	59058.627701	1387.0	0.000082	59071.042587	1410.5	0.000075
59046.739114	1364.5	0.000086	59058.891008	1387.5	0.000076	59071.306873	1411.0	0.000055
59047.003837	1365.0	0.000072	59059.156332	1388.0	0.000087	59075.269512	1418.5	0.000073
59047.267302	1365.5	0.000078	59059.419938	1388.5	0.000098	59075.534009	1419.0	0.000081
59047.532523	1366.0	0.000084	59059.684201	1389.0	0.000085	59075.797714	1419.5	0.000062
59047.795808	1366.5	0.000088	59059.947683	1389.5	0.000082	59076.061853	1420.0	0.000062
59048.060856	1367.0	0.000103	59060.212760	1390.0	0.000092	59076.326308	1420.5	0.000064
59049.380015	1369.5	0.000075	59060.475958	1390.5	0.000095	59076.590032	1421.0	0.000062
59049.645507	1370.0	0.000053	59062.059702	1393.5	0.000070	59076.853883	1421.5	0.000055
59049.908823	1370.5	0.000083	59062.324947	1394.0	0.000050	59077.118281	1422.0	0.000062
59050.174129	1371.0	0.000064	59062.588783	1394.5	0.000071	59077.382065	1422.5	0.000067
59050.436919	1371.5	0.000082	59062.853569	1395.0	0.000059	59077.646390	1423.0	0.000059
59050.701683	1372.0	0.000067	59063.117122	1395.5	0.000067	59077.910784	1423.5	0.000066
59050.964997	1372.5	0.000087	59063.381628	1396.0	0.000061	59078.174122	1424.0	0.000060
59051.230902	1373.0	0.000069	59063.645362	1396.5	0.000089	59078.438383	1424.5	0.000070
59051.494492	1373.5	0.000072	59063.910690	1397.0	0.000048	59078.702977	1425.0	0.000068
59051.759397	1374.0	0.000084	59064.174613	1397.5	0.000075	59078.966824	1425.5	0.000072
59052.023123	1374.5	0.000066	59064.438258	1398.0	0.000061	59079.231510	1426.0	0.000067
59052.286863	1375.0	0.000092	59064.701572	1398.5	0.000073	59079.494570	1426.5	0.000069
59052.550981	1375.5	0.000093	59064.966169	1399.0	0.000058	59079.759833	1427.0	0.000052
59052.815849	1376.0	0.000072	59065.229515	1399.5	0.000056	59080.023589	1427.5	0.000084
59053.079435	1376.5	0.000067	59065.494779	1400.0	0.000064	59080.288840	1428.0	0.000071
59053.343861	1377.0	0.000073	59065.758511	1400.5	0.000068	59080.550777	1428.5	0.000060
59053.607858	1377.5	0.000073	59066.022631	1401.0	0.000049	59080.816015	1429.0	0.000081
59053.871578	1378.0	0.000075	59066.286854	1401.5	0.000080	59081.078603	1429.5	0.000104
59054.135531	1378.5	0.000086	59066.550834	1402.0	0.000066	59081.343803	1430.0	0.000056
59054.400856	1379.0	0.000070	59066.814401	1402.5	0.000071	59081.607927	1430.5	0.000072
59054.664783	1379.5	0.000076	59067.079190	1403.0	0.000075	59081.872450	1431.0	0.000053
59054.929447	1380.0	0.000116	59067.343567	1403.5	0.000082	59082.135381	1431.5	0.000083
59055.192947	1380.5	0.000070	59067.608362	1404.0	0.000055	59082.401060	1432.0	0.000060
59055.457163	1381.0	0.000070	59067.872493	1404.5	0.000062	59082.663637	1432.5	0.000064
59055.721318	1381.5	0.000065	59068.136789	1405.0	0.000065	59082.928962	1433.0	0.000067
59055.985884	1382.0	0.000059	59068.399954	1405.5	0.000074	59083.192191	1433.5	0.000088
59056.250243	1382.5	0.000079	59068.665304	1406.0	0.000058	59083.457546	1434.0	0.000064
59056.514623	1383.0	0.000070	59068.928573	1406.5	0.000051	59083.720955	1434.5	0.000100
59056.778127	1383.5	0.000087	59069.193428	1407.0	0.000064	59083.985845	1435.0	0.000051
59057.043001	1384.0	0.000055	59069.456733	1407.5	0.000063	59084.249018	1435.5	0.000080
59057.306395	1384.5	0.000079	59069.721349	1408.0	0.000069	59084.513792	1436.0	0.000065
59057.570656	1385.0	0.000075	59069.984985	1408.5	0.000060			
59057.834665	1385.5	0.000084	59070.250315	1409.0	0.000075			

stars from the broadening profile. The RV points obtained in this way are tabulated in Table 6.

The phase-folded RV points (after the correcting for the orbital motion around the center of mass of the whole quadruple system) together with the best-fit photodynamical solution (see below, in Sect. 3) are plotted in Fig. 7.

3   JOINT ANALYSIS OF THE AVAILABLE DATA

We used the software package LIGHTCURVEFACTORY (see Borkovits et al. 2019a, 2020, and further references therein) to carry out a complex spectro-photodynamical modeling of the system based on the data collected in Sect. 2. LIGHTCURVEFACTORY calculates stellar positions and velocities for each object and emulates the light-, RV and ETV

curves of any arbitrary quadruple system (having either 2+2 or 2+1+1 hierarchies), including mutual eclipses amongst any two (or more) components. Moreover, the software may (optionally) use built-in, pre-calculated PARSEC isochrone tables[7] (Bressan et al. 2012) to constrain the stellar parameters theoretically through their evolution tracks, and also to model the combined stellar energy distribution (SED) of the four stars. To solve the inverse problem, the code employs a Markov Chain Monte Carlo (MCMC) based parameter search with an implementation of the generic Metropolis-Hastings algorithm (see e.g. Ford 2005).

Our combined analysis is primarily based on the following observational inputs: (*i*) the high-quality *TESS*

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[7] These tables generated via the web based tool CMD 3.3, http://stev.oapd.inaf.it/cgi-bin/cmd

Table 4. Eclipse times of BG Ind binary A determined from archival, ground-based photometric measurements.

BJD −2 400 000	Cycle no.	std. dev. (d)	BJD −2 400 000	Cycle no.	std. dev. (d)	BJD −2 400 000	Cycle no.	std. dev. (d)
45905.747600[a]	-8483.5	0.001000	56203.239825	-1450.0	0.000915	56574.380595	-1196.5	0.000149
46670.720391[b]	-7961.0	0.001007	56205.439280	-1448.5	0.000161	56577.309551	-1194.5	0.000229
46695.610663[b]	-7944.0	0.001082	56211.296461	-1444.5	0.000117	56585.361049	-1189.0	0.000328
50749.603297[c]	-5175.0	0.000200	56213.491330	-1443.0	0.000201	56599.268820	-1179.5	0.000453
50752.535697[c]	-5173.0	0.001000	56219.346595	-1439.0	0.000305	56601.468138	-1178.0	0.000489
50760.584197[c]	-5167.5	0.000200	56241.309142	-1424.0	0.000095	56604.395227	-1176.0	0.000115
51096.587705[c]	-4938.0	0.000200	56246.434520	-1420.5	0.000777	56607.321861	-1174.0	0.000190
51137.582236[c]	-4910.0	0.000200	56249.362519	-1418.5	0.000130	56615.374282	-1168.5	0.000208
55783.058458[d]	-1737.0	0.000800	56252.292900	-1416.5	0.000312	56819.609731	-1029.0	0.000231
55837.954960[d]	-1699.5	0.000200	56257.416438	-1413.0	0.000460	56822.538145	-1027.0	0.000935
56114.664682	-1510.5	0.000263	56260.343245	-1411.0	0.000176	56868.660034	-995.5	0.000387
56120.518954	-1506.5	0.000139	56450.673609	-1281.0	0.000270	56871.584259	-993.5	0.000166
56125.644943	-1503.0	0.000088	56453.601199	-1279.0	0.000060	56888.423844	-982.0	0.000073
56128.573285	-1501.0	0.000546	56464.581939	-1271.5	0.000063	56893.547481	-978.5	0.000130
56134.428810	-1497.0	0.000679	56480.684869	-1260.5	0.000345	56902.333131	-972.5	0.000358
56139.552054	-1493.5	0.000130	56483.612309	-1258.5	0.000133	56904.529609	-971.0	0.000119
56150.534786	-1486.0	0.000459	56491.664083	-1253.0	0.000358	56907.456133	-969.0	0.000216
56153.464094	-1484.0	0.000251	56500.449720	-1247.0	0.000224	56913.311888	-965.0	0.000214
56158.587494	-1480.5	0.000071	56505.572608	-1243.5	0.000104	56923.563562	-958.0	0.000627
56161.515700	-1478.5	0.000094	56508.501528	-1241.5	0.000094	56926.489918	-956.0	0.000094
56166.641916	-1475.0	0.000295	56524.605414	-1230.5	0.000128	56934.541525	-950.5	0.000231
56167.373624	-1474.5	0.000103	56538.511594	-1221.0	0.000128	56935.274575	-950.0	0.000141
56175.425003	-1469.0	0.000060	56541.440387	-1219.0	0.000123	56945.521544	-943.0	0.000392
56177.620729	-1467.5	0.025388	56547.296407	-1215.0	0.000896	56948.450675	-941.0	0.000107
56178.352518	-1467.0	0.000088	56557.543911	-1208.0	0.000203	56951.379494	-939.0	0.000137
56186.407006	-1461.5	0.000160	56558.279472	-1207.5	0.000209	56954.308374	-937.0	0.000156
56188.603921	-1460.0	0.000307	56560.472694	-1206.0	0.000120	56967.484429	-928.0	0.000531
56191.527603	-1458.0	0.000357	56565.594480	-1202.5	0.001205	56970.412294	-926.0	0.000236
56197.385670	-1454.0	0.000289	56566.330280	-1202.0	0.000089			
56202.511063	-1450.5	0.000278	56569.257382	-1200.0	0.000281			

Notes. Integer and half-integer cycle numbers refer to primary and secondary eclipses, respectively. Times of minima between cycle nos. −1737.0 and −926.0 were determined from WASP measurements. The sources of the few other, older eclipse times are follows: (a) van Hamme & Manfroid (1988); (b) present paper, determined from the ESO archival timeseries, see Sect. 2.4; (c) present paper, determined from unpublished observations of Jens Viggo Clausen; (d) present paper, observations of co-author MB.
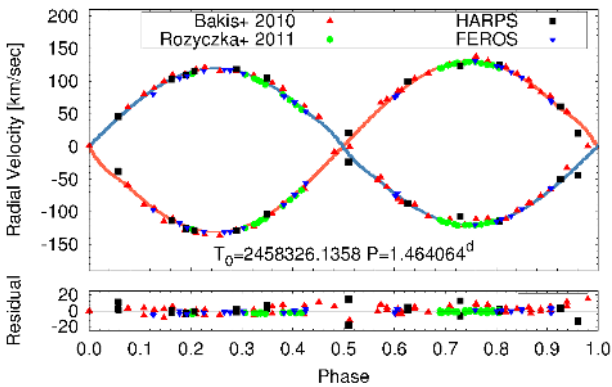

Figure 7. Phase-folded RV curve of the brighter binary, A, in BG Ind after the removal of the contribution of the orbital motion around the center of mass of the quadruple system, and the systemic RV (γ) of the quadruple, as well. (The values to be removed were calculated from the best-fit joint spectro-photodynamical model, see below in Sect. 3). The origin of each set of data points are noted in the key box. Red and blue lines stand for the model solutions for the photometric primary and secondary (spectroscopic secondary and primary) components, respectively.

lightcurve (see Sect. 2.2), (*ii*) the RV data available in the literature (see Sect. 2.7), (*iii*) the ETV data calculated from all the available photometric observations (see Sect. 2.6 and Tables 2–4) and, (*iv*) the observed passband magnitudes of the target taken from standard catalogs (see Sect. 2.1 and Table 1).

Note that, in its present form, LIGHTCURVEFACTORY is unable to handle period variations caused by non-few-body perturbations. Therefore, we did not model the small linear period variation of binary A which manifests itself in the form of quadratic deviations of the first few ETV points (see above, in Sect. 2.6.3). Considering the fact that our analysis depends primarily on the *TESS* measurements obtained over the last 2.5 years, and on the radial velocity data gathered within a relatively narrow eight-year-long interval about a decade ago, we do not expect that such a small, long-term effect will have any significant influence on our results. Nevertheless, we will return to this question in Sect. 4.

Regarding the archival photometric observations, we decided not to use the photometric fits to these lightcurves themselves for the analysis. This decision was based mainly on the fact that their large scatter was found to be of the same order as, or even higher than, the eclipse depths of the

Table 5. Light-travel-time orbital solution for BG Ind A from a classical ETV analysis of its outer orbit

Parameter	value
$P_{\rm out}$ [d]	721.2 ± 0.9
$a_{\rm AB} \sin i_{\rm out}$ [$R_\odot$]	161 ± 9
$e_{\rm out}$	0.21 ± 0.06
$\omega_{\rm out}$ [°]	345 ± 8
$\tau_{\rm out}$ [BJD]	$2\,458\,678 \pm 16$
$f(m_{\rm B})$ [$M_\odot$]	0.11 ± 0.02
$K_{\rm A}$ [km s^{-1}]	11.5 ± 0.6
$\dot{P}_{\rm A}/P_{\rm A}$ [10^{-8} yr^{-1}]	6.1 ± 0.4

Notes. $a_{\rm AB} \sin i_{\rm out}$ denotes the line-of-sight projected semi-major axis of the outer orbit of binary A around the center of mass of the quadruple system, while the other orbital elements and associated parameters are noted in their usual manner. Moreover, we tabulate two derived parameters: $f(m_{\rm B})$, and $K_{\rm A}$ which are the mass function and the amplitude of the radial velocity curve of the center of mass of binary A on its outer orbit. Finally, in the last row, we give the rate of the continuous period variation of binary A, which is derived from eqn. (1).

Table 6. Unpublished ESO archive RV data for BG Ind.

BJD −2 400 000	$V_{\rm Aa}$	$\sigma_{\rm Aa}$	$V_{\rm Ab}$	$\sigma_{\rm Ab}$	Instr.
		(km s^{-1})			
53191.746239	104.898	2.107	17.981	2.731	HARPS
53191.747289	103.736	2.558	16.724	2.676	HARPS
53196.615059	134.576	2.113	-30.806	1.888	FEROS
53196.616399	133.978	2.006	-30.566	1.895	FEROS
53205.719319	-16.673	2.060	133.718	1.792	FEROS
53205.720649	-17.291	2.040	134.728	1.842	FEROS
55468.519001	161.283	2.291	-54.899	2.114	HARPS
55471.513061	173.835	2.285	-70.150	1.922	HARPS
55477.491091	175.936	2.223	-70.213	1.905	HARPS
55478.473551	14.546	2.592	85.063	1.895	HARPS
55479.598661	-48.246	1.202	181.327	2.374	HARPS
55479.711481	-55.896	2.364	183.279	1.943	HARPS
55535.522132	5.406	2.841	117.879	2.621	HARPS
56449.951656	14.266	1.994	60.955	1.400	HARPS
56450.949046	148.986	2.185	-85.547	1.943	HARPS
56473.803176	-77.384	2.044	160.731	1.827	FEROS
56473.844686	-69.309	2.033	151.984	1.870	FEROS
56475.840366	145.259	2.001	-87.001	1.924	FEROS
56475.964466	153.088	2.046	-93.602	1.869	FEROS
56553.665337	138.067	2.543	-70.208	2.350	HARPS
56577.496218	-51.494	2.467	133.971	1.974	HARPS
56906.616809	112.284	2.517	2.106	2.109	FEROS
56908.568369	-60.329	1.962	189.225	1.840	FEROS
56908.644099	-52.057	2.098	180.060	1.812	FEROS
56908.702749	-37.210	1.950	164.030	1.893	FEROS
56908.784749	-5.251	2.000	129.609	1.942	FEROS
56908.787889	-5.126	1.842	127.316	1.864	FEROS
56909.512639	125.816	2.033	-15.650	1.842	FEROS
56909.808189	-13.712	1.987	137.056	1.815	FEROS
56910.570019	138.822	1.930	-38.326	1.623	FEROS
56910.645749	163.451	2.081	-58.828	1.893	FEROS
56910.715679	174.478	2.007	-71.379	1.874	FEROS
56910.774629	175.871	2.132	-72.403	1.874	FEROS
56910.829279	169.785	2.038	-65.114	1.846	FEROS

faint binary B. We did, however, utilize in our analysis the most relevant information that could be mined from these observations, namely the best of the mid-eclipse times that could be derived from these data. Furthermore, the other benefit of these data is that they reveal that the eclipse depths of binary A have remained constant during the last ~ 40 years, the relevance of which will be discussed later.

Before the analysis, we also took further preparatory steps on the *TESS* lightcurve. In order to save computational time we binned the 2-min cadence data into 30-min bins, and in the following analyses we worked with the binned data. While carrying out some preliminary fitting runs on this 30-min binned Sector 1 lightcurve, we realized that the residual curve exhibits small amplitude, quasiperiodic variations. We had also found very similar patterns in the residual lightcurve at the end of the prior lightcurve disentangling process (Sect. 2.5), i.e., after removing both binaries from the original time series. We therefore concluded that these small quasi-cyclic variations cannot be the consequence of some misadjustments of the lightcurve parameters during our analysis, but should be real effects. In order to find the dominant frequencies of these fluctuations we calculated the power spectrum of the residual curve with a discrete Fourier-transform. We found two independent sets of frequency peaks. The frequencies of one set were close to the orbital frequency of binary A (and its multiples), while the other set was clearly related to the orbital frequency of binary B.

Similar to what we have done in some of our previous work (see e.g. Borkovits et al. 2018) we modelled these fluctuations during our analysis in the following manner. We found that the use of the two most dominant frequencies of both sets of frequencies resulted in a significant improvement in the solutions. The process itself works as follows: in each trial step, after the removal of the blended EB lightcurves from the observed data, the residual lightcurve is modelled with harmonic functions of the four fixed frequencies, of which the eight (plus one) coefficients are obtained via matrix inversion. Then, this mathematical model of the residual lightcurve is added to the double binary model lightcurve and the actual χ^2 value is calculated for this mixed model lightcurve. Finally note, since the fluctuations were found to be quasi-periodic instead of strictly periodic, we found that our process is the most effective if we use only a short section of the *TESS* lightcurve. Therefore, for the main portion of our analysis we used only a seven-day-long section of the Sector 1 *TESS* lightcurve.[8]

The combined analyses were carried out in two different stages. In the first stage we worked only with astrophysical model-independent parameters. Therefore, we fitted simultaneously only the *TESS* lightcurve, and the RV and ETV curves, but did not include SED data and theoretical stellar isochrones. During this phase the 20 adjusted parameters were as follows:

(i) Seven lightcurve related parameters: the temperature ratios of $(T_2/T_1)_{\rm A,B}$ and $T_{\rm Ba}/T_{Aa}$, i.e. the secondary over

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[8] We will show later, however, that by arbitrarily choosing another section of the lightcurve, we obtain very similar results, well within the 1σ statistical uncertainties of most of the adjusted parameters.

primary temperature ratios of both binaries, and the ratio of the temperatures of the two primaries; the durations of the two primary eclipses $(\Delta t_{\mathrm{pri}})_{\mathrm{A,B}}$; the ratios of the radii in both pairs $(R_2/R_1)_{\mathrm{A,B}}$; and the gravity darkening coefficients of the two stars of binary A $(\beta_{\mathrm{Aa,Ab}})$.

(ii) One parameter for each inner binary orbit, i.e. the observed inclinations $i_{\mathrm{A,B}}$ of the orbital planes of binary A and binary B, and five orbital parameters of the outer orbit: period (P_{out}), time of periastron passage τ_{out}, eccentricity and argument of periastron $(e\cos\omega)_{\mathrm{out}}$ and $(e\sin\omega)_{\mathrm{out}}$, and the inclination i_{out}.

(iii) Four mass-related parameters: the masses of the two primaries $(m_{\mathrm{Aa,Ba}})$, and the mass ratios of the two binaries $(q_{\mathrm{A,B}})$.

Regarding the other orbital parameters of the inner binaries, the periods $(P_{\mathrm{A,B}})$ of these EBs, as well as their orbital phase (through the time of an arbitrary primary eclipse – $\mathcal{T}_{\mathrm{A,B}}^{\mathrm{pri}}$) were constrained internally through the ETV data. Furthermore, the eccentricities of both inner orbits were set to zero. Moreover, for the large $P_{\mathrm{out}}/P_{\mathrm{A,B}}$ ratios we found that all three orbits (two inner binary orbits and the outer orbit) can be considered as pure, unperturbed Keplerian motion. Due to this latter consideration our dataset does not contain any information about the positions of the orbital nodes relative to each other. Therefore, the sixth orbital element, the longitude of the node of each orbit $(\Omega_{\mathrm{A,B;out}})$ was fixed at zero. Finally, we note that the systemic radial velocity of the whole quadruple system (γ) was also constrained internally by minimizing the χ^2_{RV} contribution a posteriori in each trial step.

Turning to the atmospheric parameters of the four stars, in contrast to our previous analyses, we now adjust the gravity darkening coefficients (β) of the strongly non-spheroidal components of the bright binary A. The reason is that, in contrast to the widely used classic model of Lucy (1967) which predicts a unique gravity darkening coefficient of $\beta = 0.32$ for all convective stars, recently Claret & Bloemen (2011) have shown that the true relations are much more complicated. This is especially true for stars close to the transition region between convective and radiative envelopes, where the components of binary A are located. On the other hand, in the case of binary B, we kept fixed the usual value of $\beta = 0.32$ prescribed in Lucy's model. Other atmospheric parameters, such as the logarithmic limb-darkening coefficients $(x,y)_{\mathrm{TESS}}$ were interpolated in each trial step with the use of passband-dependent tables downloaded from the Phoebe 1.0 Legacy page[9]. These tables are based on Castelli & Kurucz (2004) atmospheric models and are primarily used for the original version of the Phoebe software (Prša & Zwitter 2005). Furthermore, for the components of the bright binary A $(Aa$ and $Ab)$, we include the reflection/irradiation effect into the lightcurve model and, therefore, we take into account the bolometric limb-darkening coefficients $(x,y)_{\mathrm{bol}}$, interpolating them in each trial step in the same manner as was done with the passband-dependent coefficients.

At this stage of the analysis we required only one further parameter that is undetermined by the model and, therefore, has to be set externally. This was the effective temperature of the primary of binary A, which was set (and kept fixed) according to the findings of Rozyczka et al. (2011).

At the end of this stage of the analysis we obtained accurate dynamical masses not only for the two members of the bright binary A but, in addition, we obtained the total dynamical mass of binary B.[10] Furthermore, the temperature ratio of the two primaries provides reliable information about the characteristics of the two stars forming the faint binary B. Finally at this stage, the orbital elements of the three orbits were also accurately determined.

In the next and final stage of the analysis, we included the SED information into the analysis as well as the built-in PARSEC tables. Now the seven lightcurve-related parameters described above were no longer adjusted but, instead, the radii and temperatures of all the four stars were constrained, i.e., recalculated at the beginning of each trial step by interpolating their values from the three-dimensional (mass, metallicity, age) grids of the PARSEC tables[11]. During this phase of the analysis three additional adjustable quantities were introduced, including (i) the metallicity $([M/H])$ and (ii) the (logarithmic) age of the quadruple. These two parameters, together with the mass of the given components, determined the position of each star within the PARSEC grids and, therefore, determine the interpolated fundamental stellar parameters and theoretical passband magnitudes. The third parameter (iii) was the stellar extinction $E(B-V)$ for the SED fitting. Moreover, while fitting the model SED to the dereddened observed SED points, the distance of the system comes in as an additional parameter. The software constrains this parameter a posteriori in each trial step by minimizing the value of χ^2_{SED}.

After some initial trials, however, we found it necessary to introduce a fourth, extra parameter to adjust, namely the age of the evolved component of binary A, in order to obtain model lightcurves which yield similarly low χ^2_{LC} values to the ones obtained in the previous astrophysical model-independent stage. This procedure requires some further explanation. It is generally expected that the components of a close binary (multiple) system are coeval. Theories, however, allow for small departures from exact coevality (see, e.g. Tokovinin 2018b), which during some critical rapid stages of stellar evolution might be significant. Furthermore, even in the case of exact coevality, the approximative nature of our interpolation method certainly carries with it inherent inaccuracies which might lead to modest discrepancies in the derived stellar parameters, especially during the very rapid sensitive evolutionary stage of the evolved star in binary A. Therefore, as a counterbalance to these uncertainties, we allowed for the age of the evolved component to be set independently from the other three stars.

In Table 7 we tabulate the median values and the 1σ statistical uncertainties of the parameters obtained during the last stage of our analysis. The synthetic model light curves derived from the best-fit joint solution are displayed in Figs. 1 and 2. The corresponding ETV and RV curves

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[9] http://phoebe-project.org/1.0/download

[10] This latter was driven mainly by the amplitude ratios of the cyclic ETV curves of the two binaries and also by the varying systemic RV of binary A (not to be confused with the systemic RV of the whole quadruple, γ, described above).

[11] The interpolation method was described in detail in Borkovits et al. (2020)

are presented in Figs. 6 and 7, respectively. Finally, in the two panels of Fig. 8, we illustrate the SED-fitting part of the combined solution both in the flux and the passband magnitude domain.

4   RESULTS AND DISCUSSION

4.1   Orbital configuration

Our analysis confirms the hierarchical 2+2 type quadruple star nature of BG Ind. Thanks to the available high-quality *TESS* photometry and the long-term ground-based photometric and spectroscopic observations, BG Ind now takes its place as (i) one of the most compact 2+2 quadruples known, as well as (ii) the quadruple system with the most accurately known stellar masses and other stellar parameters. The outer period of the system is found to be $P_{\mathrm{out}} = 721 \pm 3$ days, which is the shortest amongst doubly eclipsing quadruple systems with an accurately known outer period.[12] Note, however, that despite the relatively short outer period, both outer to inner period ratios are large enough ($P_{\mathrm{out}}/P_{\mathrm{A}} \approx 492$, $P_{\mathrm{out}}/P_{\mathrm{B}} \approx 1365$) so that we do not expect readily measurable short-term mutual three- (four-) body perturbations. In other words, all three orbits can be considered as essentially purely Keplerian. The outer orbit is moderately eccentric with $e_{\mathrm{out}} = 0.21 \pm 0.05$, and is seen nearly along the direction of the minor axis ($\omega_{\mathrm{out}} = 2° \pm 9°$).

These relatively small uncertainties in the BG Ind quadruple system, however, should be treated with some caution. The two main reasons for this caveat can nicely be seen in the ETV plots in Figure 6. First, due to the unlucky fact that the outer period is nearly exactly equal to two years ($P_{\mathrm{out}} \approx 1.973\,\mathrm{yr}$), the annual observing seasons of the target can, and do, miss the most informative two parts of the ETV curve, i.e. its two extrema. Second, as was discussed above, the very first eight pre-WASP ETV points show clear deviations from the pure LTTE solution, and might indicate a continuous, constant increase in the orbital period of binary A, i.e., $\dot{P}_{\mathrm{A}}$.

Though the modelling of $\dot{P}_{\mathrm{A}}$ was not included in the comprehensive spectro-photodynamical approach, its effect can be quantified by comparing the orbital parameters of the outer orbit obtained through the classic, analytic LTTE+quadratic solution of the ETV of binary A (Table 5) with the detailed spectro-photodynamical model (Table 7). As one can see, the outer period and eccentricity match well within their estimated uncertainties, while the argument of pericenter, the periastron passage time and the RV amplitudes are discrepant at the $2 - 3\sigma$ level. Therefore, we can conclude that, as was expected, the omission of the quadratic ETV term in the complex spectro-photodynamical analysis did not influence our basic solutions, but suggests that the

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[12] We emphasize that this holds only for *doubly eclipsing* 2+2 quadruples. The shortest period known 2+2 quadruple system, VW LMi has a much shorter outer period of $P_{\mathrm{out}} = 355$ days (Pribulla et al. 2008, 2020). Furthermore, in the case of the doubly eclipsing quadruple star EPIC 220204960 Rappaport et al. (2017) reported that the outer period is very likely between 300 and 500 d, but an accurate value for that system is unknown.

actual uncertainties in the orbital elements should be somewhat larger than cited in Table 7.

The inclination of the outer orbit is found to be $i_{\mathrm{out}} = 86° \pm 5°$. On the other hand, the inclinations of the two inner eclipsing binaries are found to be $i_{\mathrm{A}} = 73°.1 \pm 0°.1$ and $i_{\mathrm{B}} = 84°.3 \pm 0°.9$. From these values, and in the absence any information on the longitude of the nodes of the three orbits ($\Omega_{\mathrm{A,B,out}}$), the only thing one can say is that the whole quadruple system is certainly not perfectly flat. Since the mutual inclination of two planes cannot be smaller than the difference between the two observed inclinations of the planes considered (and, cannot be larger than their sum), the inclination of the bright binary A relative to the outer orbit must surely exceed $\approx 13 \pm 5°$, but may even reach $90°$. (Similarly, the mutual inclination between the orbital plane of binary B and the outer orbital plane may be anywhere between coplanar and perpendicular.)

As a consequence of such misalignments, one may expect the binary's orbital plane to precess. In that case, eclipse depth variations should be observed, or even the disappearance of the eclipses on a longer time-scale. The period of forced precession of a binary orbital plane in a hierarchical triple system can be well approximated with the expression (see e.g. Söderhjelm 1975)

$$(P_{\mathrm{prec}})_{\mathrm{A}} = \frac{4}{3} \frac{1 + q_{\mathrm{out}}}{q_{\mathrm{out}}} \frac{P_{\mathrm{out}}^2}{P_{\mathrm{A}}} \left(1 - e_{\mathrm{out}}^2\right)^{3/2} \left[\frac{C}{G_2} \cos(i_{\mathrm{m}})_{\mathrm{A}}\right]^{-1},$$ (2)

where C represents the total orbital angular momentum of the quadruple, while G_2 is the orbital angular momentum stored in the outer orbit. In the present situation it can be readily seen that $C/G_2 \approx 1$, i.e., the majority of the orbital angular momentum of the quadruple is stored in the outer orbit. Therefore, with $q_{\mathrm{out}} \equiv M_B/M_A \simeq 0.5$ one can easily show that $(P_{\mathrm{prec}})_{\mathrm{A}} \gtrsim 4200\,\mathrm{yr}$. From this we conclude that there is no chance of detecting eclipse-depth variations given the available span of the observations.

4.2   Astrophysical properties and evolutionary status of the four stars

Turning to the fundamental astrophysical parameters of the bright components of binary A, we compare our results to those of the former thorough analysis of Rozyczka et al. (2011). The masses of the stars found in the two analyses agree quite well: 1.432 ± 0.020 vs. 1.428 ± 0.008 for the more evolved component, and 1.315 ± 0.025 vs. 1.293 ± 0.008 for its less evolved companion, where the first of each pair are from the current work. This agreement is good to $\lesssim 1\,\sigma$, in units of our error bars. The uncertainties given in Rozyczka et al. (2011) are smaller than ours by factors of 2-3. One should keep in mind, however, that Rozyczka et al. (2011) estimated their uncertainties from an analysis of their own set of RV data which have smaller rms residuals than the combined set of RVs that we used. Furthermore, during their final RV analysis they corrected the RV values for the distortions of the stellar components with the use of the Wilson–Devinney code (Wilson & Devinney 1971; Wilson 1979). In contrast to this, in our study, the effects of the distortions of the stars on the RV data are automatically taken into account within

Table 7. Orbital and astrophysical parameters of BG Ind from the joint photodynamical lightcurve, RV, ETV, SED and `PARSEC` isochrone solution.

orbital elements			
	subsystem		
	A	B	A–B
P_{a} [days]	$1.464065^{+0.000002}_{-0.000002}$	$0.528349^{+0.000002}_{-0.000002}$	$720.9^{+3.4}_{-3.1}$
a [R$_\odot$]	$7.602^{+0.038}_{-0.043}$	$3.025^{+0.011}_{-0.016}$	$540.4^{+2.7}_{-2.2}$
e	0	0	$0.209^{+0.028}_{-0.048}$
ω [deg]	–	–	$1.6^{+9.2}_{-8.8}$
i [deg]	$73.27^{+0.06}_{-0.13}$	$84.29^{+0.85}_{-0.87}$	$85.5^{+3.1}_{-6.3}$
$\mathcal{T}^{\mathrm{pri}}$ [BJD - 2 400 000]	$58326.1362^{+0.0011}_{-0.0012}$	$58325.8072^{+0.0025}_{-0.0022}$	
τ [BJD - 2400000]	–	–	58699^{+14}_{-21}
mass ratio [$q = m_{\mathrm{sec}}/m_{\mathrm{pri}}$]	$0.919^{+0.010}_{-0.006}$	$0.932^{+0.014}_{-0.015}$	$0.483^{+0.007}_{-0.005}$
K_{pri} [km s^{-1}]	$120.47^{+1.12}_{-0.75}$	$138.98^{+1.23}_{-1.39}$	$12.57^{+0.17}_{-0.24}$
K_{sec} [km s^{-1}]	$130.99^{+0.48}_{-0.50}$	$149.25^{+1.09}_{-1.13}$	$26.02^{+0.31}_{-0.50}$
γ [km s^{-1}]		–	$48.69^{+0.29}_{-0.59}$

stellar parameters			
	Aa	Ab	Ba	Bb

Relative quantities and atmospheric properties			
	Aa	Ab	Ba	Bb
fractional radius^b [R/a]	$0.3084^{+0.0016}_{-0.0044}$	$0.2096^{+0.0052}_{-0.0019}$	$0.2120^{+0.0013}_{-0.0014}$	$0.2019^{+0.0018}_{-0.0020}$
fractional flux [in *TESS*-band]	$0.6294^{+0.0064}_{-0.0153}$	$0.3475^{+0.0165}_{-0.0060}$	$0.0133^{+0.0009}_{-0.0008}$	$0.0092^{+0.0008}_{-0.0007}$
x^c_{bol}	$0.676^{+0.001}_{-0.001}$	$0.671^{+0.001}_{-0.001}$	...	...
y^c_{bol}	$0.174^{+0.002}_{-0.003}$	$0.198^{+0.002}_{-0.003}$	...	...
x^c_{TESS}	$0.631^{+0.001}_{-0.001}$	$0.618^{+0.001}_{-0.001}$	$0.720^{+0.001}_{-0.001}$	$0.724^{+0.001}_{-0.002}$
y^c_{TESS}	$0.347^{+0.002}_{-0.002}$	$0.354^{+0.002}_{-0.003}$	$0.281^{+0.005}_{-0.005}$	$0.325^{+0.019}_{-0.018}$
β^d	$0.11^{+0.10}_{-0.07}$	$0.64^{+0.29}_{-0.33}$	0.32	0.32

Physical Quantities			
m [M$_\odot$]	$1.432^{+0.015}_{-0.024}$	$1.315^{+0.026}_{-0.023}$	$0.688^{+0.008}_{-0.011}$	$0.640^{+0.010}_{-0.011}$
R^b [R$_\odot$]	$2.339^{+0.016}_{-0.021}$	$1.592^{+0.047}_{-0.019}$	$0.642^{+0.005}_{-0.007}$	$0.611^{+0.008}_{-0.009}$
T^b_{eff} [K]	6442^{+29}_{-28}	6816^{+26}_{-26}	4609^{+48}_{-49}	4327^{+62}_{-57}
L^b_{bol} [L$_\odot$]	$8.433^{+0.199}_{-0.169}$	$4.934^{+0.279}_{-0.179}$	$0.167^{+0.009}_{-0.009}$	$0.118^{+0.009}_{-0.009}$
M^b_{bol}	$2.45^{+0.02}_{-0.03}$	$3.04^{+0.04}_{-0.06}$	$6.72^{+0.06}_{-0.06}$	$7.09^{+0.08}_{-0.08}$
M^b_V	$2.45^{+0.03}_{-0.03}$	$3.02^{+0.04}_{-0.06}$	$7.23^{+0.10}_{-0.09}$	$7.83^{+0.13}_{-0.13}$
$\log g^b$ [dex]	$3.852^{+0.011}_{-0.005}$	$4.150^{+0.007}_{-0.016}$	$4.660^{+0.003}_{-0.002}$	$4.672^{+0.004}_{-0.004}$

Global Quantities		
$\log(\mathrm{age})^{b,e}$ [dex]	$9.40^{+0.02}_{-0.01}$	$9.32^{+0.03}_{-0.01}$
$[M/H]^b$ [dex]		$-0.189^{+0.038}_{-0.037}$
$E(B-V)$ [mag]		$0.018^{+0.015}_{-0.013}$
$(M_V)^b_{\mathrm{tot}}$		$1.93^{+0.03}_{-0.03}$
distance [pc]		$69.7^{+0.7}_{-0.9}$

Notes. a: Time of the inferior conjunction of the secondary component (i.e. mid-time of a primary eeclipse); *b*: Interpolated (or derived) from the `PARSEC` isochrones; *c*: interpolated linear (x) and logarithmic (y) limb-darkening coefficients. Note, bolometric coefficients used only during the calculation of the reflection effect, therefore they were not set for binary B; *d*: gravity darkening coefficients; *e*: the age of the evolved primary component of binary A was allowed to vary independently of the other three stars – see text for details.

Lightcurvfactory. Therefore, we consider our somewhat larger uncertainties to be more realistic.

The radii of the two stars in binary A exhibit slightly larger differences between the two studies. Our analysis has yielded $R_{\mathrm{Aa}} = 2.34 \pm 0.02\,\mathrm{R}_\odot$ for the evolved component, while Rozyczka et al. (2011) obtained the somewhat smaller value of $2.29 \pm 0.02\,\mathrm{R}_\odot$.[13] For the other less evolved star we

found $R_{\mathrm{Ab}} = 1.59 \pm 0.04\,\mathrm{R}_\odot$, in contrast to their somewhat larger value of $1.68 \pm 0.04\,\mathrm{R}_\odot$. Keeping in mind, however, that the sum of the (fractional) radii of the two stars (i.e. $R_{\mathrm{Aa}}/a_A + R_{\mathrm{Ab}}/a_A$) is one of the most robustly determined parameters of an eclipsing binary's lightcurve (though, it

[13] Note, however, that they analysed six different lightcurves sep-

arately, and their results were scattered between $2.17\,\mathrm{R}_\odot$ and $2.40\,\mathrm{R}_\odot$.
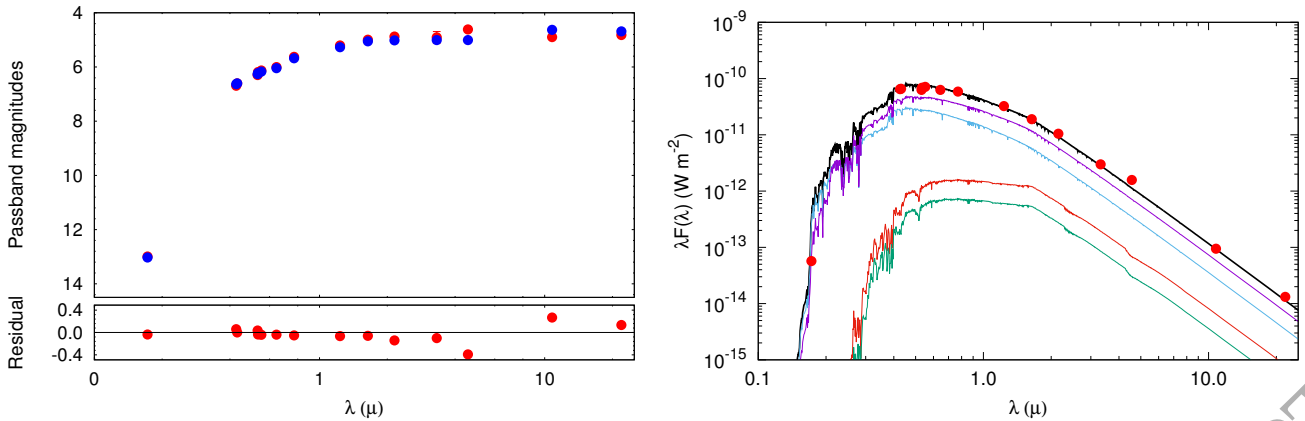
Figure 8. The summed SED of the four stars of BG Ind both in the magnitude and the flux domains. The *left panel* displays the cataloged values of the passband magnitudes (red filled circles; tabulated in Table 1) versus the model passband magnitudes derived from the absolute passband magnitudes interpolated with the use of the PARSEC tables (blue filled circles). In the *right panel* the dereddened observed magnitudes are converted into the flux domain (red filled circles), and overplotted with the quasi-continuous summed SED for the quadruple star system (thick black line). This SED is computed from the Castelli & Kurucz (2004) ATLAS9 stellar atmospheres models (http://wwwuser.oats.inaf.it/castelli/grids/gridp00k2odfnew/fp00k2tab.html). The separate SEDs of the four stars are also shown with thin green, black and purple lines, respectively.

does depend sensitively on the inclination), one can easily check, that this sum agrees well in the two solutions.[14]

   The relatively small discrepancies in the individual radii between our analysis and that of Rozyczka et al. (2011), can likely be explained by some combination of the following three effects. First, prior studies did not consider the small extra flux contribution (of $\approx 2.2\%$) coming from binary B, and the lightcurve distortions caused by its eclipses. Second, we used the high quality *TESS* photometry whose superiority over the former ground-based measurements is unquestionable. Third, we allowed for the reflection/irradiation effect which made our analysis more realistic, but this effect was not considered during the previous analyses. In conclusion, we emphasize again that the discrepancy in radii is fairly small.

   We also found small departures in the effective temperatures of the two components of binary A compared with the previous results. Our results, which hinge to a large degree on the fit of the combined 4-star SED, resulted in slightly larger temperatures. We found $T_{\rm Aa} = 6442 \pm 29\,{\rm K}$ and $T_{\rm Ab} = 6816 \pm 26\,{\rm K}$ in contrast to $6350 \pm 260\,{\rm K}$ and $6650 \pm 230\,{\rm K}$ (Rozyczka et al. 2011). Note, however, that our results are within the uncertainties of Rozyczka et al. (2011). On the other hand, by using the temperatures given by Rozyczka et al. (2011), Stassun & Torres (2016) found a consistent SED solution for the binary. Of course, they did not consider the contribution of binary B, which might give

a small excess at the red wing of the SED, and in turn which might force the fit toward slightly lower temperatures.

   There is, however, an even more significant discrepancy between the system distance inferred from the SED solution and the trigonometric distance deduced from Gaia's measurements. Our solution has resulted in a photometric distance of $d = 69.7 \pm 0.8\,{\rm pc}$, while Bailer-Jones et al. (2018) using the Gaia DR2 measurements have obtained $d_{\rm DR2} = 51.0 \pm 0.5\,{\rm pc}$.[15] The situation is more complicated than this seemingly straightforward discrepancy. First, the trigonometric distance that can be calculated from the new reduction of HIPPARCOS parallaxes (van Leeuwen 2007) is $d_{\rm HIP} = 67.1^{+2.8}_{-2.2}$, which is within 1σ of our result. Furthermore, Stassun & Torres (2016) used a similar SED modeling analysis to ours, and found a photometric distance of $d_{\rm Stassun+16} = 66.7\,{\rm pc}$. Again, this is much closer to our result and that of HIPPARCOS than to the Gaia distance. However, since we know the fundamental stellar parameters of the dominant A binary quite accurately (including the bolometric luminosities) independent of the distance, and others have used partially different methods[16] to find a very similar distance, we tentatively conclude that the published Gaia DR2 and EDR3 parallaxes are probably subject to some systematic error. This discrepancy might have arisen from the fact that the period of the outer orbit in BG Ind is very close to two years ($P_{\rm out} = 1.973\,{\rm yr}$) and, further-

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[14] Note also that due to the significant tidal and rotational oblateness of the two stars, they will no longer be spherical; therefore, it should be clarified what is meant by 'radius'. We cite the volume equivalent radius and assume that Rozyczka et al. (2011) used the same definition. On the other hand, we note that in the above mentioned relation for the sum of the fractional radii, the so-called 'side' radius, (i.e. measured in the star's equatorial plane, in the direction perpendicular to the line joining the two stars) should be the relevant one during eclipses. The volume equivalent and side radii for our stars, however, agree to better than 1%.

[15] The parallaxes published in Gaia DR2 and EDR3 are well within 1σ of each other, and therefore we can assume that the distance derived from the EDR3 data will not differ significantly from the published DR2 distance.

[16] Rozyczka et al. (2011) have determined the temperatures with a combination of spectroscopic analysis and lightcurve fitting. Stassun & Torres (2016) utilized SED modelling. And our study included a combination of lightcurve and SED fitting, where the consistency of the obtained stellar parameters were also probed by modelling the RV curves and identifying appropriate PARSEC isochrones tracks.

more, the semi-major axis of BG Ind A's ellipse around the center of mass of the quadruple system is $a_{\mathrm{out,A}} \approx 0.82\,\mathrm{AU}$. Therefore, the combination of the orbital motion of the photocenter of binary A along the outer orbit and a period near 2 years, may be responsible for causing some problems with Gaia's trigonometric parallax determination.

Turning to the newly discovered binary B, we find it to be a pair of two mid K-type dwarfs with a near unit mass ratio of $q_{\mathrm{B}} = 0.93\pm0.01$. Despite the fact that the flux contribution from this binary is only about 2% in the *TESS* photometric band, due to the high quality *TESS* photometry of this rather bright quadruple, we were able to obtain quite good lightcurves (see Fig. 3) and a robust dynamical model for binary B. The out-of-eclipse sections of the lightcurve of binary B are distorted, and we explain that by chromospheric activity, i.e. stellar spots, which are quite usual for stars with thick convective envelopes. As one can see in the middle panel of Fig. 3, for Sector 1 data these variations can be well modelled (mathematically) with two Fourier-terms having frequencies equal to the orbital frequency and its first harmonic. On the other hand, in the case of Year 3 (i.e. Sectors 27 and 28 data) the same Fourier-representation was found to be less satisfactory, and we therefore added two additional harmonics of the binary B orbital frequency to the Fourier-representation (see the lowest panel of Fig. 3). However, even in this case, one can still notice some imperfections in the fit. We explain this fact with a possible rapid variation in the chromospheric activity which induces brightness fluctuations that cannot be well represented by a few smooth harmonics (even after averaging the lightcurve over a few weeks of the *TESS* observations). Note, however, that this discrepancy is less than $\sim$100 ppm, and therefore, it would remain under the detection limit for any ground-based photometric observations.

The timings of the shallow eclipses from binary B are in accord with the 1.973-yr-periodicity in the ETV curve found from binary A, both of which are dominated by the light travel time effect (Fig. 6). Moreover, the dynamically deduced total mass of binary B, coupled with the dimensionless stellar parameters, lead to physical parameters of the stars in binary B. And these, according to the PARSEC tables we used, are fully consistent with the parameters of two main-sequence K-dwarfs, having the same age and metallicity as those of the members of the bright binary A. Therefore, there is no question that the two eclipsing binaries form a compact, gravitationally bound, hierarchical quadruple star system.

Regarding the global parameters of the quadruple, the combined solution prefers a slightly metal-deficient abundance of $[M/H] = -0.19 \pm 0.04$ which, again, is in perfect agreement with the previous result of $[Fe/H] = -0.2 \pm 0.1$ by Rozyczka et al. (2011). As mentioned before, we did not enforce strict coevality among the stellar components during our analyses and found an age of $\tau_{\mathrm{Aa}} = 2.51 \pm 0.12\,\mathrm{Gyr}$ for the evolved primary, and $\tau_{\mathrm{Ab,B}} = 2.14 \pm 0.10\,\mathrm{Gyr}$ for the three main-sequence components. These two ages differ by $\simeq 370 \pm 150\,\mathrm{Myr}$, or $\sim$7% of the age, with a significance of only $2.5\,\sigma$. We consider this discrepancy to be not a 'small' departure from the coevality (though it does not have a high statistical significance). Our impression is that it might arise from the rapid rotation as well as the tidal distortion of the evolved star. Therefore, it is possible that the spherical stellar radius given by the PARSEC tables would not strictly equal the volume-equivalent radius of a strongly tidally distorted star.

Finally, we briefly discuss the question of the likely continuous period increase in the binary A period, detected through the systematic deviations of the very first ETV points from a simple linearly sloped LTTE model. Such period variations have been observed in a fair number of EBs. In the case of semi-detached and contact systems, the most common explanation is some kind of mass exchange between the stellar components. However, since all the previous studies have found that BG Ind A is a detached system, and our detailed analysis confirms this scenario, this period increase cannot be explained via mass transfer. (And this is not to mention the fact that, in this case, the increasing period would imply that the less evolved, lower mass secondary star would be the mass-donor, which is an unphysical scenario.) On the other hand, however, as was shown, e.g. by Pringle (1985) and Demircan et al. (2006) mass loss from a close binary star (e.g., due to a stellar wind) always leads to an increasing period. Therefore, the quadratic ETV-term in BG Ind A might imply an enhanced stellar wind from the surface of the evolved component. For a quantitative study of this possibility, further observations of high-quality eclipse times over a longer time interval would be extremely useful.

We also note that naturally, an LTTE effect driven by a more distant, low-mass fifth stellar component may also be the correct explanation. Obviously, the confirmation or refutation of this scenario also requires further eclipse follow up observations.

5 SUMMARY AND CONCLUSIONS

In this paper we report the discovery of the doubly eclipsing quadruple nature of the previously known, bright, southern eclipsing binary BG Ind. We present the first comprehensive analysis of BG Ind in its entirety. *TESS* observations provided high-precision photometry covering two intervals of one and two months, respectively, and separated by two years. Even though these high-quality *TESS* observations covered only short segments of the outer orbit with $P_{\mathrm{out}} = 1.973\,\mathrm{yr}$, we were able to use ground-based archival lightcurve and RV data to determine accurately the orbital and dynamical parameters of the system.

BG Ind is found to have one of the shortest outer periods among all quadruple systems having a 2+2 hierarchy. According to the recent version of the Multiple Stellar Catalog (Tokovinin 2018a) there are only five such systems (including BG Ind) with outer periods shorter than 3 years. These are tabulated in Table 8.

The remarkably small number of such compact 2+2 quadruples that are known probably arises from an observational selection effect rather than for astrophysical reasons—specifically, they are quite difficult to discover. In contrast with the discovery of a third companion of a known binary star, which can be made by, e.g., astrometry, long-term RV measurements (or, in the case of an EB) ETV studies, or in exceptional cases, observing serendipitous extra eclipses, the binary nature of such a third component would remain hidden in most cases. The only rare exceptions are when the

Table 8. List of hierarchical 2+2 quadruple systems with $P_{out} <$ 3 yr.

Identifier	P_{out}	P_A	P_B	References
VW LMi	355	0.478	7.931*	1
EPIC 220204960	300–500	13.274	14.416	2
BG Ind	721	1.464	0.528	3
TIC 278956474	858	5.488	5.674	4
V994 Her	1063	2.083	1.420	5

Notes. * The 7.931-day-period binary in VW LMi does not exhibit eclipses.
References: (1) Pribulla et al. (2008); (2) Rappaport et al. (2017); (3) present paper; (4) Rowden et al. (2020); (5) Zasche & Uhlař (2016)

second binary happens to be also an EB[17] (as is the case in four of the five systems listed in Table 8), or it is bright enough to be observable as a second spectroscopic binary (as in the case of the fifth tabulated system, VW LMi). Furthermore, another possibility for discovering the binarity of a component in a binary or multiple star system might arise from its faintness in comparison to its dynamically determined mass (as in the cases of κ For, Tokovinin 2013 and ζ Cnc C, Tokovinin 2017).[18]

BG Ind is also one of only a very few compact quadruple systems where the key parameters of all fours stars are known with an accuracy of better than $\sim$3%, including masses, radii, and T_{eff} values. Likewise, the three sets of orbital parameters, such as periods, semi-major axes, eccentricities, and inclination angles are all known rather precisely. The one notable exception is that we do not have any way of determining the mutual inclination angles among the three orbital planes. The observational inclination angles are all close to edge on (i.e., $90°$), and we might surmise from statistical arguments that the most likely configuration is nearly coplanar for all three orbits. But we cannot be certain that this is indeed the case.

Future interferometric and astrometric observations may help to solve this problem. The semimajor axis of the outer orbit is $\simeq$ 36 mas, so in principle it is resolvable by speckle interferometry or adaptive optics. However, the high contrast between the A and B binaries makes resolution with single-dish telescopes a challenging task. A much bet-

ter prospect is offered by long-baseline interferometers, e.g. the GRAVITY instrument at VLTI[19] (Gravity collaboration 2017). The contrast in the K band is more favorable compared to the visible ($\Delta K = 2.98$ mag vs $\Delta V = 4.73$ mag), so the visibility and phase modulation caused by the outer pair can be well measured. Furthermore, the spectral resolution of $R \sim 4000$ offered by GRAVITY will allow detection of opposite phase shifts in the spectral lines of Aa and Ab at times near their maximum separation, thus enabling one to measure the orientation of the Aa,Ab orbit (its semimajor axis is 0.5 mas) and, perhaps, even the orbit of Ba,Bb.

The orientation of the outer orbit on the sky will also be known from future Gaia data releases because the amplitude of the photocentric orbit is quite large, $\sim$12 mas. The proper motion anomaly (difference between the short-term proper motion measured by Gaia and the long-term proper motion deduced from the HIPPARCOS and Gaia positions) is quite large, $(+19.4, -13.7)$ mas/yr (Brandt 2018). Moreover, the long-term proper motion deduced from the HIPPARCOS and Gaia positions, $(+5.4, +29.0)$ mas/yr, is close to $(+5.0, +30.2)$ mas/yr measured by Gaia EDR3 on a 2-yr time base that effectively averages the outer orbit, while Gaia DR2 proper motion measured on a 1.5 yr baseline is substantially different.

Finally since BG Ind is so bright, and the eclipses of binary A are relatively deep (at $\sim$15%), we encourage amateurs to continue the eclipse timing. The historical archival data are very helpful in this regard, but not as accurate as targeted observations of this star would be with small to modest sized telescopes. Furthermore, a secure verification of the suspected continuous period change of binary A also needs long-term, continuous follow up timing observations.

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[17] On the other hand, despite the fact that such large photometric surveys such as, e.g., the gound-based Optical Gravitational Lensing Experiment (OGLE; Udalski et al. 2015), or the *TESS* mission have observed hundreds of lightcurves exhibiting blends of at least two EBs, the gravitationally bound nature of the blended EBs have been proven definitively for only a relatively small fraction of these objects (see, e. g. Zasche et al. 2019).

[18] In the case of the compact hierarchical triple system IU Aur Drechsel et al. (1994) and Özdemir et al. (2003) have also concluded that the large third mass vs. small third light and weak spectroscopic signal discrepancies could be resolved by postulating that the third companion is a binary itself. If this assumption was true, IU Aur would be the shortest outer period 2+2 quadruple with $P_{out} = 294$ d, but the system needs further investigations. Moreover, note that most recently Marcadon et al. (2020) have proposed that the $P_{out} = 180.4$ days period outer component of V1200 Cen might also be a binary, forming a more tight half-year-long period quadruple system.

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DATA AVAILABILITY

The *TESS* data underlying this article were accessed from MAST (Barbara A. Mikulski Archive for Space Telescopes) Portal (https://mast.stsci.edu/portal/Mashup/Clients/Mast/Portal.html). A part of the data were derived from sources in public domain as given in the respective footnotes. The derived data generated in this research and the code used for the photodynamical analysis will be shared on reasonable request to the corresponding author.

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