



Dark Energy Survey Year 1 Results: The Photometric Data Set for Cosmology

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Abstract

We describe the creation, content, and validation of the Dark Energy Survey (DES) internal year-one cosmology data set, Y1A1 GOLD, in support of upcoming cosmological analyses. The Y1A1 GOLD data set is assembled from multiple epochs of DES imaging and consists of calibrated photometric zero-points, object catalogs, and ancillary data products—e.g., maps of survey depth and observing conditions, star–galaxy classification, and photometric redshift estimates—that are necessary for accurate cosmological analyses. The Y1A1 GOLD wide-area object catalog consists of ~ 137 million objects detected in co-added images covering ~ 1800 deg² in the DES *grizY* filters. The 10σ limiting magnitude for galaxies is $g = 23.4$, $r = 23.2$, $i = 22.5$, $z = 21.8$, and $Y = 20.1$. Photometric calibration of Y1A1 GOLD was performed by combining nightly zero-point solutions with stellar locus regression, and the absolute calibration accuracy is better than 2% over the survey area. DES Y1A1 GOLD is the largest photometric data set at the achieved depth to date, enabling precise measurements of cosmic acceleration at $z \lesssim 1$.

Key words: catalogs – cosmology: observations – surveys – techniques: image processing – techniques: photometric

1. Introduction

The Dark Energy Survey (DES; DES Collaboration 2005, 2016) is a photometric survey utilizing the Dark Energy Camera (DECam; Flaugher et al. 2015) on the Blanco 4 m telescope at Cerro Tololo Inter-American Observatory (CTIO) in Chile to observe ~ 5000 deg² of the southern sky in five broadband filters, g , r , i , z , Y , ranging from ~ 400 to ~ 1060 nm (Li et al. 2016; Burke et al. 2017).⁵¹ The primary goal of DES is to study the origin of cosmic acceleration and the nature of dark energy through four key probes: weak lensing, large-scale structure, galaxy clusters, and Type Ia supernovae (SNe). More generally, DES provides a rich scientific data set and has already had a significant impact beyond cosmology (e.g., DES Collaboration 2016).

Precision measurements of dark energy with DES rely on an unprecedented survey data set and a comprehensive understanding of the survey performance. It is necessary to identify, characterize, and mitigate the influences of variable observing conditions, data processing artifacts, photometric calibration nonuniformity, and astrophysical foregrounds. For example, photometric calibration must be accurate and uniform to avoid introducing noise and bias into photometric redshift estimates. Studies of galaxy clustering depend on a detailed knowledge of survey coverage, galaxy detection efficiency, and the accuracy of recovered galaxy properties. Furthermore, detailed modeling of the point-spread function (PSF) and instrument response is required to perform galaxy shape measurements on objects that are fainter than the detection limit of a single DES image. The scale and complexity of assembling, characterizing, and validating the DES data motivate a collaborative effort that draws upon and enables a wide range of scientific analyses.

Here we describe the creation, composition, and validation of the DES first-year (Y1) data set in support of cosmological analyses (shown schematically in Figure 1). While this data set is currently proprietary to the DES Collaboration, this document is intended to serve as a reference for these data products when they become publicly available.⁵² Observing for DES Y1 spanned from 2013 August to 2014 February and covered $\sim 40\%$ of the DES footprint, averaging three to four visits per band. The resulting images were processed through

the DES data management (DESDM) system (Ngeow et al. 2006; Mohr et al. 2008; Sevilla et al. 2011; Desai et al. 2012; Mohr et al. 2012; Morganson et al. 2018) and assembled into the DES year-one annual data set (Y1A1). Y1A1 consists of reduced single-epoch images and object catalogs (known as “Y1A1 FINALCUT”), along with multiepoch co-added images and associated multiband catalogs (known as “Y1A1 COADD”). Photometric calibration of Y1A1 was performed globally on a CCD-to-CCD basis, and maps of the survey coverage and depth were assembled with the *mangle* software suite (Hamilton & Tegmark 2004; Swanson et al. 2008). The Y1A1 data set covers ~ 2000 deg² in any single filter with inhomogeneous coverage and depth. In total, ~ 1800 deg² of the Y1A1 footprint has simultaneous coverage in all five DES filters.

The desired precision of DES cosmological analyses motivates further refinement of Y1A1. The resulting data set, referred to as Y1A1 GOLD, is accompanied by extensive validation and ancillary data products to facilitate cosmological analyses. The primary components of Y1A1 GOLD are (Figure 1) (1) a multiband photometric object catalog subselected from the Y1A1 COADD object catalog; (2) an adjusted photometric calibration to improve uniformity over the survey footprint; (3) shape and photometry information from a simultaneous multiepoch, multioject fit; (4) a set of ancillary maps quantifying survey characteristics using the HEALPix rasterization scheme (Górski et al. 2005); and (5) several value-added quantities for high-level analyses (i.e., a star–galaxy classifier and photo- z estimates). When creating the Y1A1 GOLD object catalog, several classes of nonphysical, spurious, or otherwise problematic objects were identified and either flagged or removed. The calibrated magnitudes of objects were also corrected for interstellar extinction using a stellar locus regression (SLR) technique. The ancillary data products associated with Y1A1 GOLD accurately quantify the characteristics of the survey, further mitigating the impact of systematic uncertainties. A high-level summary of the performance of Y1A1 GOLD is tabulated in Table 1.

Our purpose here is to document the production and performance of the Y1A1 GOLD data set in support of upcoming DES cosmology analyses. We start by describing the DES Y1 observations in Section 2 and briefly reviewing the image reduction pipeline applied to produce the Y1A1 data set in Section 3. In Section 4 we describe the photometric calibration of the Y1A1 data, and in Section 5 we describe the image co-addition process. In Section 6 we discuss the creation

⁵¹ The DECam filter throughput is publicly available at <http://www.ctio.noao.edu/noao/node/1033>.

⁵² Note that the DES Y1 cosmology data set described here is distinct from the forthcoming DES public data release, which will include data from the first 3 yr of DES.

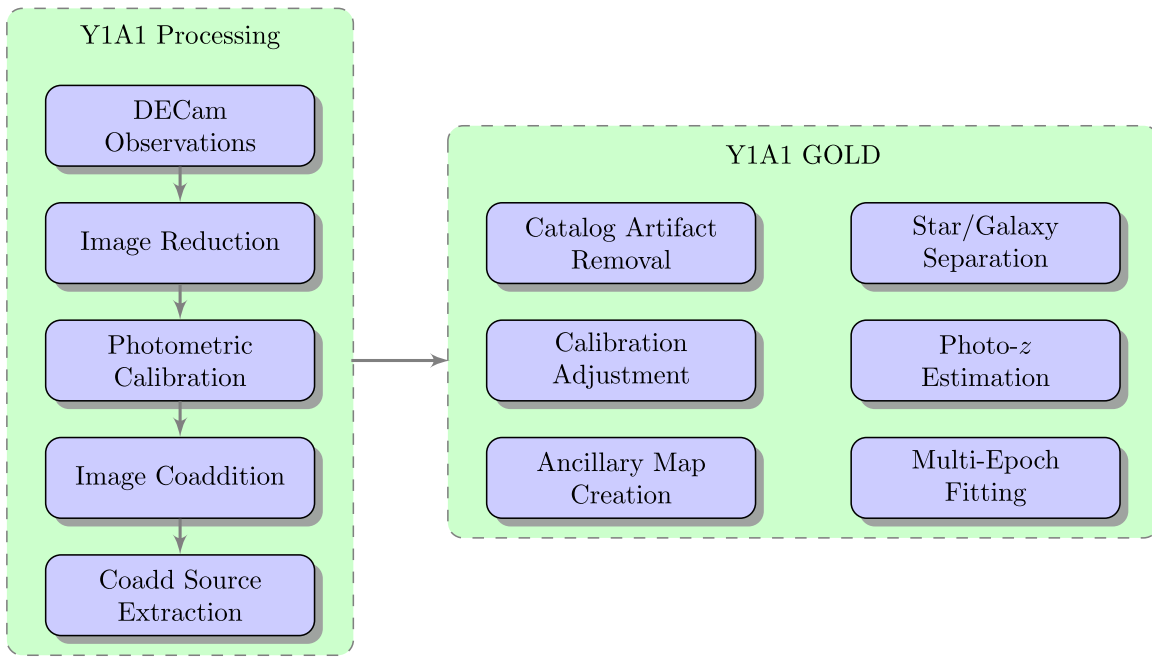


Figure 1. Schematic of the constituents of the Y1A1 processing (left) and the additional Y1A1 GOLD products (right).

of unique object catalogs, while in Sections 7 and 8 we describe the ancillary maps and value-added quantities produced to complement the Y1A1 GOLD catalog. We briefly conclude in Section 9.

2. Data Collection

DES has been allocated 105 nights per year on the Blanco telescope starting in 2013. The first year of DES observing spanned from 2013 August 31 to 2014 February 9 and consisted of both full and half nights.⁵³ Details on DES operation and data collection are provided by Diehl et al. (2014); here we briefly summarize some of the key details relevant to the creation of Y1A1 GOLD.

DES consists of two observing programs: a shallower wide-area survey and a deeper time-domain (SN) survey (Figure 2). The DES wide-area survey footprint covers $\sim 5000 \text{ deg}^2$ with 90 s exposures in *griz* and 45 s exposures in *Y*. A single imaging pass over this footprint, called a “tiling,” collects science data over $\sim 83\%$ of the survey footprint owing to inefficiencies in the pointing layout and camera footprint (e.g., area not covered owing to gaps between CCDs, nonfunctioning CCDs, and problematic area near the edges of the CCDs). The DECam pointings for each tiling are shifted relative to each other by a large fraction of the camera field of view in a dither pattern designed to maximize uniformity and distribute repeated detections of a given object over the focal plane. During Y1, DES observed $\sim 2000 \text{ deg}^2$ of the wide-area survey footprint with three to four dithered tilings per filter. The Y1 footprint consisted of two areas: one near the celestial equator including Stripe 82 (S82; Annis et al. 2014), and a much larger area that was also observed by the South Pole Telescope (SPT; Carlstrom et al. 2011). During Y1, DES collected 17,671 wide-area survey exposures in a variety of observing conditions (Diehl et al. 2014).

The SN survey observes 10 fields in four filters (*griz*) on a regular cadence to detect and characterize SNE through difference imaging (Kessler et al. 2015). Longer exposure times ($\geq 150 \text{ s}$) and frequent repeated visits result in a significantly deeper survey in the SN fields. All 10 SN fields reside within the DES wide-area footprint, but only two were covered by wide-area imaging in Y1 (Figure 2). Over the course of Y1, DES collected a total of 2699 time-domain survey exposures.

In addition to the wide-area and SN survey fields, two auxiliary fields outside the DES footprint were observed to aid in the training of photometric redshifts and star-galaxy classification. Fields overlapping with COSMOS (Scoville et al. 2007) and VVDS-14h (Le Fèvre et al. 2005) were observed during the DES Science Verification (SV) period.⁵⁴ These observations are deeper than most of the Y1 wide-area survey.

During DES operation, sets of biases and flat-field calibration exposures were taken in each filter before each night of observing. Standard-star fields were observed at three different airmasses at the beginning and end of each night unless conditions were obviously nonphotometric.⁵⁵ Cloud cover was monitored continuously during observing by the RASICAM all-sky infrared camera (Lewis et al. 2010; Reil et al. 2014).

DES images the sky whenever weather allows the Blanco dome to be open, resulting in some exposures being taken in very poor conditions. Thus, data quality monitoring is essential to select exposures that meet the scientific requirements of the survey. The quality of exposures is evaluated based on the PSF, sky brightness, and sky transparency. For each exposure, we define t_{eff} to be the ratio between the actual exposure time and the exposure time necessary to achieve the same signal-to-noise ratio (S/N) for point sources observed in nominal conditions

⁵⁴ Data from the DECam Science Verification period is available at <https://des.ncsa.illinois.edu/releases/sva1>.

⁵⁵ On DES half nights only two standard-star fields were observed at the midpoint of the night.

⁵³ Several exposures taken during engineering time earlier in 2013 August were also included in the Y1A1 data set.

Table 1
Y1A1 GOLD Data Quality Summary

Parameter	Band					Reference
	<i>g</i>	<i>r</i>	<i>i</i>	<i>z</i>	<i>Y</i>	
Median PSF FWHM	1''25	1''07	0''97	0''89	1''07	Section 7.2
Sky coverage (in all bands)	1786 deg ²				1773 deg ²	Section 7.3
Astrometric accuracy	25 mas (relative); <300 mas (external)					Section 5.1
Absolute photometric uncertainty (mmag)	14	4	2	15	32	Section 4.4
Relative photometric uniformity (mmag)	19	22	20	20	18	Section 4.4
Completeness limit (95%)	23.6	23.4	22.9	22.4		Section 6.4
Co-add galaxy magnitude limit (10 σ) ^a	23.4 ^{+0.14} _{-0.40}	23.2 ^{+0.13} _{-0.37}	22.5 ^{+0.14} _{-0.34}	21.8 ^{+0.12} _{-0.37}	20.1 ^{+0.18} _{-0.33}	Section 7.1
Multiepoch galaxy magnitude limit (10 σ) ^a	23.7 ^{+0.07} _{-0.40}	23.5 ^{+0.16} _{-0.29}	22.9 ^{+0.14} _{-0.30}	22.2 ^{+0.14} _{-0.32}	...	Section 7.1
Galaxy selection ($i \leq 22$)	Efficiency > 98%; contamination < 3%					Section 8.1
Stellar selection ($i \leq 22$)	Efficiency > 86%; contamination < 6%					Section 8.1

Note.

^a The quoted values correspond to the mode, 16th percentile, and 84th percentile of the magnitude limit distribution. Using the median instead of the mode reduces the magnitude limit by ~ 0.05 mag.

(Nielsen et al. 2015).⁵⁶ To pass preliminary data quality cuts, wide-area survey exposures must have $t_{\text{eff}} > 0.3$ in the *r*, *i*, and *z* bands and $t_{\text{eff}} > 0.2$ in the *g* and *Y* bands. The median measured t_{eff} for Y1 was $t_{\text{eff}} = 0.75$ in the *r*, *i*, and *z* bands and $t_{\text{eff}} = 0.49$ in the *g* and *Y* bands. In contrast, the preliminary data quality cuts for SN exposures require FWHM $< 2''$ and that a 20th-magnitude simulated source have S/N > 20 (> 80) for the shallow (deep) exposures (Kessler et al. 2015). Of the exposures taken during Y1, 82% of the wide-area exposures and 95% of the SN exposures passed their respective data quality criteria (i.e., did not require re-observation). A number of additional exposures were removed from the Y1 data set owing to instrumental artifacts (scattered light and internal reflections from bright stars, contaminating light from airplanes, poor telescope tracking, shutter malfunctions, dome occultations, etc.). In total, the Y1A1 FINALCUT processing consists of 16,857 DECam exposures, including wide field, SN, auxiliary fields, and standard stars.

3. Image Processing

The DESDM system is responsible for reducing, cataloging, and distributing DES data. Earlier iterations of the DESDM image processing pipeline are outlined in Sevilla et al. (2011) and Mohr et al. (2012), and a more detailed summary with updates for the forthcoming DES 3 yr processing is available in Bernstein et al. (2017a) and Morganson et al. (2018). Here we briefly summarize the single-epoch image processing steps applied during the DES Y1A1 FINALCUT campaign. The Y1A1 FINALCUT campaign resulted in ~ 20 TB of processed images and a catalog of ~ 740 million detected objects.

1. *Overscan and Cross Talk*: Each DECam CCD has two amplifiers for converting photo-carrier counts to analog-to-digital units (ADU). For each amplifier, the average in the overscan region was calculated and subtracted on a row-by-row basis. Cross talk is manifested as low-level leakage of electronic signals between different readout amplifiers and is observed at the level of $\sim 10^{-3}$ for pairs of amplifiers on the same CCD and $\sim 10^{-4}$ – 10^{-6} for pairs

of amplifiers on different CCDs on the same electronic back plane. Cross talk was corrected by applying a matrix operation to the simultaneous readout of 140 amplifiers (including the amplifiers for the eight focus and alignment CCDs). The elements of the cross-talk correction matrix were derived from the median amplifier output for each “victim” channel as a function of the “source” amplifier signal for a large number of sky images. Cross talk between the DECam CCDs is found to be nonlinear when the signal on the source amplifier exceeds its saturation level—i.e., the level at which the amplifier response becomes nonlinear (Figure 2 of Bernstein et al. 2017a). This cross-talk nonlinearity was incorporated into cross-talk correction. There is no evidence for temporal variation in the cross-talk between CCDs on year timescales, and a single cross-talk matrix was used for the Y1A1 processing.

2. *Bias Correction*: A master bias frame was constructed from the average of ~ 100 zero-second exposures taken during the pre-night calibration sequences over the course of the Y1 observing season. This master bias was subtracted from each CCD to remove any residual fixed pattern noise not incorporated by the overscan correction.
3. *Bad Pixel Masking*: Bad pixel masks were created for each CCD by identifying outliers in sets of biases and *g*-band flat-field calibration exposures. These bad pixels were masked and interpolated based on values in adjacent columns. The Y1A1 processing campaign used a single static bad pixel mask. Two CCDs have failed since commissioning and were removed from Y1 processing (Figure 3; Diehl et al. 2014). CCD61 failed on 2012 November 7, and data from this CCD were not used in Y1. CCD2 failed on 2013 November 30, and data from this CCD were only included for the early months of Y1.⁵⁷
4. *Nonlinearity Correction*: Several (~ 10) CCD amplifiers have a nonlinear response at low light levels (generally below 300 ADU pixel⁻¹). For DES, this affects the sky level in short (~ 15 s) standard-star observations and wide-survey dark-sky *g*-band observations (90 s). For most other filters/exposure times, the night sky alone is

⁵⁶ The effective exposure time is defined as $T_{\text{eff}} = t_{\text{eff}} T_{\text{exp}}$, where T_{exp} is the shutter-open exposure time.

⁵⁷ CCD2 subsequently recovered on 2016 December 29.

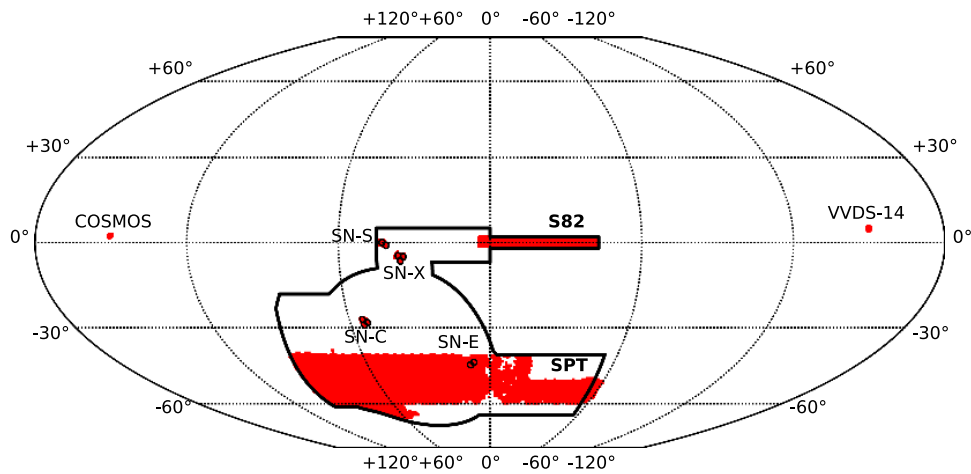


Figure 2. DES Y1A1 GOLD sky coverage in celestial coordinates (red) plotted in McBryde–Thomas flat polar quartic projection. Specific regions of the Y1A1 GOLD footprint, including the SN and auxiliary fields, are explicitly labeled (see Section 2). The nominal DES 5 yr footprint is outlined in black.

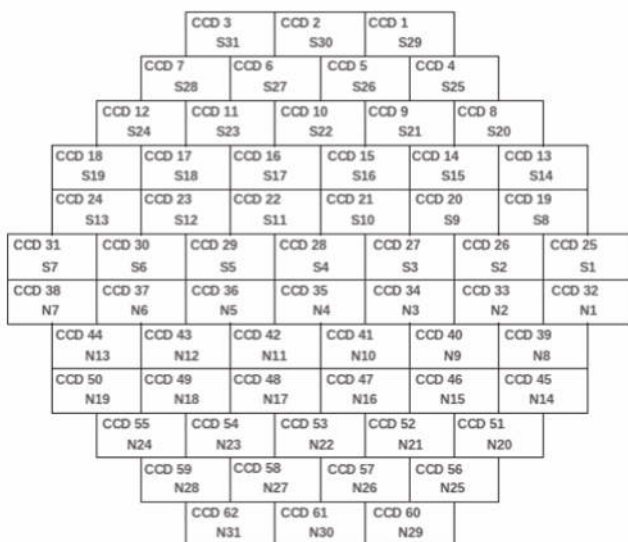
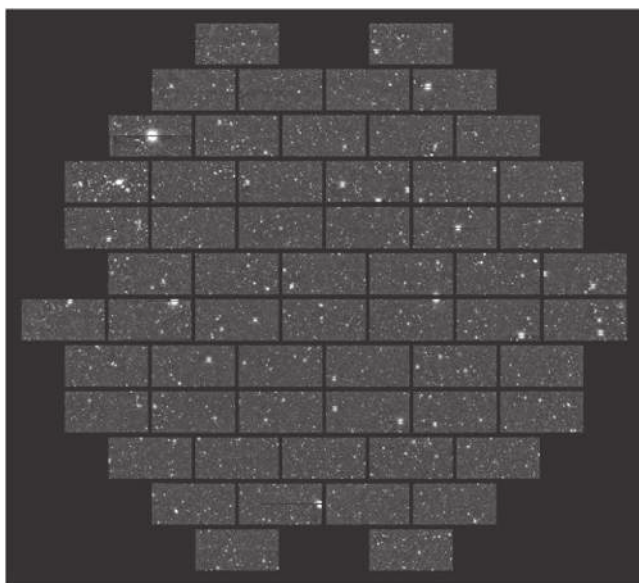


Figure 3. Processed DECam image from Y1A1 (top) and CCD layout (bottom). The three empty slots in the DECam image correspond to CCD2, CCD31, and CCD61. CCD61 failed during SV, while CCD2 failed part way through Y1. One amplifier of CCD31 has time-variable low-light-level nonlinearity, and this CCD was not processed for Y1A1.

enough to give a sufficient number of counts per pixel to make the nonlinearity correction negligible. The nonlinearity effect can be several percent at very low light levels. At very high light levels ($>2 \times 10^4$ ADU pixel⁻¹), there is also a small nonlinear behavior ($\lesssim 2\%$).

We corrected for nonlinearity at both low and high light levels using a fixed look-up table derived from calibration exposures obtained during the SV period. One amplifier on CCD31 has a time-variable nonlinear gain at the 20% level, and this CCD was excluded from Y1 processing.⁵⁸ The Y1A1 data processing did not correct for charge-induced pixel shifts (i.e., the “brighter-fatter” effect; Antilogus et al. 2014; Gruen et al. 2015), although corrections have been incorporated into more recent reductions of the DES data (Bernstein et al. 2017a).

5. *Pupil Correction:* An additive correction was applied for pupil ghosting in each exposure. As part of this process, “star flats” were created in each filter and CCD by taking multiple dithered exposures of a dense stellar field and fitting a cubic polynomial to variations in the observed brightnesses of stars. The pupil ghost correction was constructed on a CCD-by-CCD basis for each exposure from the star flat and the level of sky background (including scattered light and the night-sky pupil image). The pupil correction was scaled and subtracted from each CCD individually. This technique can leave gradients of several percent in the sky background level (worst in z and Y bands), which propagate into the reduced science images and are corrected during photometric calibration (Appendix A).⁵⁹
6. *Flat-fielding:* The response of DECam to the night sky is more stable than nightly variations in the illumination of the flat-field screen taken during pre-night calibrations. Therefore, in Y1A1 we created a single average flat-field frame for each filter from ~ 100 individual flat-field exposures. The science exposures were divided by the average flat-field frames normalized on a CCD-by-CCD basis. The pupil and flat-field corrections used for Y1A1

⁵⁸ The other amplifier on CCD31 is stable and has been recovered in more recent processing.

⁵⁹ More recent implementations of the data processing pipeline fit the additive correction over the full focal plane rather than CCD by CCD (Morganson et al. 2018).

processing remove small-scale fluctuations in the background due to pixel-size variations, i.e., tree rings, edge brightening, and tape bumps (Plazas et al. 2014). However, this correction is approximate and results in photometric measurement residuals at the level of $\sim 0.5\%$. A more rigorous correction has been applied in subsequent DES data reductions (Bernstein et al. 2017a; Morganson et al. 2018).

7. *Weight Plane Creation*: A weight plane image was created containing the inverse variance of the flat-fielded image value in each pixel. The variance estimate summed the expected Poisson noise and read noise. Saturated pixels were flagged and set to zero in the weight plane. The weight plane is used to assign relative weights to images during the co-addition process.
8. *Fringe Frame Correction*: Fringing is visible in z - and Y -band exposures. The fringing pattern is nearly identical in these bands but has a larger amplitude in the Y band. A set of templates was constructed from a stack of ~ 120 z - and Y -band exposures from DES SV. These template images were median filtered and averaged on a pixel-by-pixel level to construct a fringe frame. In the reduction pipeline, each CCD of the z - and Y -band exposures had its median sky level measured, and this sky level was used to scale the fringe frame, which was then subtracted on a CCD-by-CCD basis. The scaling method was identical to that used to scale and subtract the pupil pattern. The vast majority of exposures have a fringe residual that is $< 0.1\%$ of the sky background level. Exposures taken under the brightest conditions accepted for Y -band observing can have a fringe residual that is $\sim 0.4\%$ of the sky background level.
9. *Illumination Correction*: Light reflected from the flat-field screen fills the telescope pupil differently than the focused light of distant stars. To account for pixel-level differences in the throughput of the flat-field images, we applied a multiplicative correction to the DECam response based on the star flats. After dividing by the star flats the residual difference in response between CCDs is typically $< 2\%$ peak to peak (Appendix A.1).
10. *Preliminary Astrometric Solution*: A world coordinate system (WCS) was installed in the image header at the time of observation using a fixed distortion map derived from the star flats and an optical axis read from the telescope encoders. The pointing of each image was updated matching the centers of bright stars measured with SExtractor (Bertin & Arnouts 1996; Bertin et al. 2002) to the UCAC-4 catalog using SCAMP (Bertin 2006). This WCS was replaced by a superior one during the co-addition step (Section 5), and the astrometric accuracy of the Y1A1 GOLD catalog is described in Section 5.1.
11. *Artifact Removal*: Bright stars ($\lesssim 16$ mag) saturate the 90 s DES exposures in *griz*. Saturated pixels are set to zero in the image weight map plane. Brighter stars can produce charge overflow into pixels in the CCD readout direction. These overflow pixels are flagged in the mask plane, zeroed in the weight plane, and interpolated in the image plane. In addition, corresponding pixels on the victim amplifier of the CCD are masked owing to large nonlinear cross talk. Extremely bright oversaturated stars can leave a secondary charge overflow in the readout

register of the amplifier, conventionally called “edge bleeds” (see Figure 5 in Bernstein et al. 2017a). Edge bleeds can be located some distance from the bright star and are strongest in the rows near the readout register. These rows are identified and masked.

Energy deposited from cosmic-ray interactions with the CCDs were detected on single images using the `findCosmicRays` algorithm adopted from the LSST software stack.⁶⁰ The cosmic-ray pixels were flagged in the mask plane, zeroed in the weight plane, and interpolated in the image plane.

Long streaks produced by rapidly moving objects (i.e., meteors and Earth-orbiting satellites) were detected using a Hough transform algorithm and were also masked (Melchior et al. 2016).

12. *Single-epoch Catalog Creation*: Object catalogs were produced for each CCD using the `AstrOmatic` package (Bertin 2006). Photometric fluxes were derived using `PSFEx` and `SExtractor` for fixed apertures, the PSF model, and a galaxy model. The local sky background on each CCD was estimated by `SExtractor`. The single-epoch Y1A1 FINALCUT catalogs served as an input into the photometric calibration described in Section 4.

4. Photometric Calibration

The photometric calibration of Y1A1 was a multistep process largely following the procedure of Tucker et al. (2007). Photometric calibration was performed on the single-epoch Y1A1 FINALCUT images first on a nightly and then on a global basis. An additional calibration adjustment was derived from the stellar locus and applied at catalog level. Below we briefly describe the steps in the photometric calibration of Y1A1; a more detailed discussion of the Y1A1 photometric calibration can be found in Appendix A.

4.1. Nightly Photometric Calibration

A preliminary photometric calibration of the Y1A1 data was performed on a nightly basis. Standard-star fields were observed at various airmasses at the beginning and end of each night. These images were reduced, and the centroids of stars were matched to a set of primary standard stars from Sloan Digital Sky Survey (SDSS) DR9 (Smith et al. 2002). The DES secondary standards were then transformed to an initial DES AB photometric system via a set of transformation equations derived from SDSS DR9 and supplemented by UKIDSS DR6 (Appendix A.4). This tied the DES flux calibration of the secondary standards to SDSS and to the AB magnitude system (i.e., Padmanabhan et al. 2008).

The transformed nightly standards were used to fit a set of nightly photometric equations to model the spatial and temporal dependence of the DECam instrument throughput (Equations (6)–(10)). These equations track the accumulation of dust on the Blanco primary mirror, the relative throughput of the atmosphere at CTIO, and variations in the throughput and shape of the filter response at the location of each CCD. The nightly photometric equations produce an initial photometric calibration for all exposures taken on photometric nights. The relative calibration scatter for the nightly solution on a typical

⁶⁰ https://lsst-web.ncsa.illinois.edu/doxygen/x_masterDoxyDoc/namespacelsst_1_1meas_1_1algorithms.html

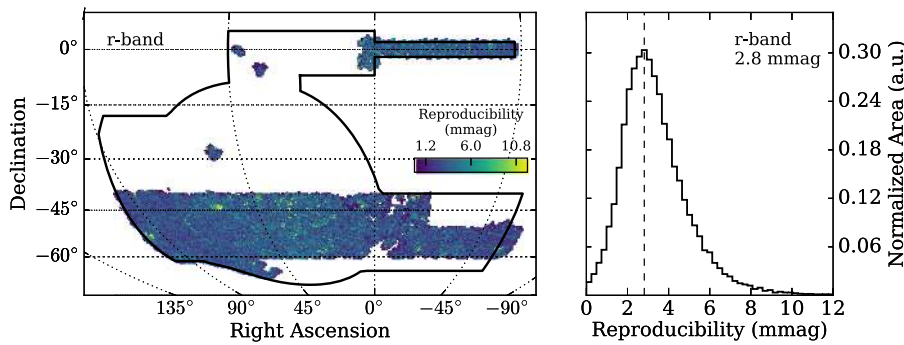


Figure 4. Internal reproducibility uncertainty for the Y1A1 *r*-band photometric zero-points calculated by comparing the rms calibrated magnitudes of stars in overlapping CCDs. The mode of the rms internal calibration uncertainty is 2.8 mmag. Similar figures for other bands are shown in Appendix A.2.

photometric night is ~ 0.02 mag rms. This nightly photometric calibration was used to anchor the relative global calibration of nonphotometric exposures described in the next section. A more detailed description can be found in Appendix A.1.

4.2. Global Calibration

We implemented a global calibration module (GCM) to derive calibrated zero-points for all exposures, including those taken under nonphotometric conditions, and to improve on the relative calibration accuracy achieved by the nightly photometric solution. The GCM procedure follows that of Glazebrook et al. (1994) and is described in more detail in Appendix A.2. Briefly, the Y1A1 data were split into regions of contiguous, overlapping images where at least one image had been previously calibrated. The calibrated images served as a reference against which other images in the grouping were calibrated. To be calibrated by the GCM, an image needed to either overlap a calibrated image or have an unbroken path of overlapping images to a calibrated image.

Following the prescription of Glazebrook et al. (1994), we estimate the rms magnitude residual for each CCD image from overlap with other CCD images. The rms distribution over all CCD images is a measure of the internal reproducibility uncertainty on small scales (the scales of overlapping CCD images) and is a measure of the precision of the overall GCM solution. We find the rms to be ~ 3 mmag (Figure 4). This uncertainty is relevant when analyzing light curves of variable objects but does not represent the internal consistency/uniformity of the relative calibrations on large scales.

While the GCM method is very precise, small systematic gradients in the flat fields of individual images can cause low-amplitude gradients over large scales. To “anchor” the fit against large-scale gradients, we used the set of nightly secondary standard stars and a sparse grid work of tertiary standard stars observed under photometric conditions and calibrated by the nightly photometric equations. The tertiary standards were chosen such that they would anchor the global solution on scales >10 – 15 deg, but on smaller scales the calibration would be dominated by the solution from overlapping uncalibrated exposures. We further examine the uniformity and absolute calibration accuracy (relative to the AB system) in Sections 4.3 and 4.4.

4.3. Photometric Calibration Adjustment

The global calibration is found to be uniform at the $\sim 2\%$ level in each band over the majority of the Y1A1 survey footprint (discussed in Section 4.4). However, nonuniformity in

the colors of objects can severely impact DES science by introducing a spatial dependence on object selection and photo- z estimation. The SLR technique uses the distinct shape of the stellar locus in color-color space to provide a relative calibration of exposures in different bands (e.g., Ivezić et al. 2004; MacDonald et al. 2004; High et al. 2009; Gilbank et al. 2011; Coupon et al. 2012; Desai et al. 2012; Kelly et al. 2014). To correct for residual spatial nonuniformity in the calibration and account for Galactic reddening (including uncertainties in the amplitude of reddening and possible variations in the effective Milky Way dust law), we have applied a secondary adjustment to the calibration of the co-add object catalogs derived from the stellar locus. Gradients in stellar population are subdominant to other calibration uncertainties in Y1A1 given the DES filter bandpasses and high Galactic latitude of the survey (e.g., High et al. 2009; Kelly et al. 2014). We followed the procedure of Drlica-Wagner et al. (2015) and applied a modified version of the `bigmacs` SLR code (Kelly et al. 2014)⁶¹ coupled with an empirical stellar locus to derive zero-point adjustments to improve the color uniformity of stars across the Y1A1 footprint. The SLR adjustment was tied to the *i*-band magnitude derived from the GCM, dereddened using the Schlegel et al. (SFD; 1998) dust map with a reddening law from O’Donnell (1994). The SLR zero-point adjustments were interpolated to the positions of each object in the catalog and were applied directly to the magnitudes of objects derived from the co-added images. In this way, the calibrated magnitudes of the Y1A1 GOLD catalog are *already corrected for interstellar extinction*. After the SLR adjustment, the color of stars was found to be uniform at the $\sim 1\%$ level across the footprint, which was verified using the red sequence of galaxies. More detail on the SLR calibration adjustment can be found in Appendix A.3.

4.4. Photometric Calibration Accuracy

To quantify the accuracy of photometric calibration, we would like to characterize the statistical distribution of $\Delta m = m_{\text{meas}} - m_{\text{true}}$, where m_{meas} and m_{true} are the measured and true magnitude of catalog objects, respectively. The characterization of the Δm distribution can be split into two components: (1) an “absolute” calibration accuracy that represents a linear shift of the Δm distribution (e.g., the mean of the distribution), and (2) a “relative” calibration accuracy that represents the spread of the Δm distribution (e.g., standard deviation of the distribution). In reality, values of m_{true} are not

⁶¹ <https://code.google.com/p/big-macs-calibrate/>

Table 2
Photometric Calibration Validation

Technique	Band				
	<i>g</i> (mmag)	<i>r</i> (mmag)	<i>i</i> (mmag)	<i>z</i> (mmag)	<i>Y</i> (mmag)
<i>Absolute Photometric Offset</i>					
GCM versus C26202	14	4	2	15	32
GCM versus Y3 FGCM	23	<1	4	11	50
<i>Relative Photometric Uniformity</i>					
GCM versus APASS/2MASS	19	22	20	20	18
GCM+SLR versus APASS/ 2MASS+SFD	25	24	20	18	15
GCM versus Y3 FGCM	14	7	8	13	15

Note. Summary of photometric calibration performance for the Y1A1 GOLD data set. See Section 4.4 for more details.

available, and we must make use of the calibrated magnitudes from other surveys or synthetic models, which have their own associated uncertainties. We describe several calibration validation studies below and summarize the results in Table 2.

The absolute calibration of the Y1A1 GOLD is tied to SDSS through the DES secondary standard stars. As an independent cross-check on the absolute photometric calibration, we examined the CALSPEC standard star, C26202 (Bohlin et al. 2014). We calculated synthetic magnitudes for C26202 by convolving *Hubble Space Telescope* (HST) spectra (stis-nic_006)⁶² with the focal-plane-averaged DECam filter throughput including atmospheric attenuation at an airmass of 1.3 (Berk et al. 1999). The predicted magnitude of C26202 in each of the DES *grizY* bands is $g = 16.695$, $r = 16.340$, $i = 16.257$, $z = 16.245$, and $Y = 16.268$. These predicted magnitudes were then compared against the pre-SLR corrected magnitudes measured by the GCM to give a “top-of-the-atmosphere” estimate of the absolute calibration uncertainty. We derive an absolute offset (in mag) of $\delta g = 0.014$, $\delta r = 0.004$, $\delta i = 0.002$, $\delta z = 0.015$, and $\delta Y = 0.032$, which we quote as the absolute photometric calibration uncertainty in Table 1.

Our primary technique for quantifying the relative photometric accuracy of Y1A1 GOLD is by comparing the calibrated magnitudes of stars against those derived from a combination of APASS (Henden & Munari 2014) and 2MASS (Skrutskie et al. 2006) (Figure 5). We perform a “top-of-the-atmosphere” comparison by calculating the difference between the GCM calibrated magnitude and the APASS/2MASS magnitude transformed to the DES system (Appendix A.4). We derive the relative calibration uncertainty as the half-width between the 16th and 84th percentiles of the difference in magnitude over the footprint: $\sigma_{68}(g) = 0.019$, $\sigma_{68}(r) = 0.022$, $\sigma_{68}(i) = 0.020$, $\sigma_{68}(z) = 0.020$, and $\sigma_{68}(Y) = 0.018$ (Table 1). These values include calibration uncertainties from both DES and APASS/2MASS and are thus a conservative upper bound on the Y1A1

GCM accuracy. We further compare the SLR-adjusted Y1A1 GOLD photometry to the transformed APASS/2MASS photometry dereddened using the SFD maps and reddening law of O’Donnell (1994). We find a dispersion of $\sigma_{68}(g) = 0.025$, $\sigma_{68}(r) = 0.024$, $\sigma_{68}(i) = 0.020$, $\sigma_{68}(z) = 0.018$, and $\sigma_{68}(Y) = 0.015$. These values include an additional contribution from differences in the reddening correction derived from the SLR and the SFD dust maps, which results in larger uncertainty in the bluer filters where interstellar reddening is more extreme. These comparisons are shown in more detail in Appendix A.5.

As an additional cross-check, we compared the “top-of-the-atmosphere” Y1A1 GCM calibration against a global calibration of the contiguous DES 3 yr data set (Y3A1). The absolute calibration of the Y3A1 data set was also tied to C26202 but made use of additional observations of this object. From comparisons against other CALSPEC standards (LDS749B and WD0308-565), the absolute calibration of Y3 is believed to be accurate at the $\sim 1\%$ level. The relative calibration of Y3 was performed over the contiguous Y3A1 footprint using an independent forward global calibration method (FGCM) and is found to be uniform at the 0.7% level (Burke et al. 2017). We checked the absolute calibration of Y1A1 GOLD by matching stars against their Y3 counterparts over the Y1A1 GOLD footprint. We found that the absolute offset between Y1A1 GCM and Y3A1 FGCM was $\delta g = 0.023$, $\delta r < 0.001$, $\delta i = 0.004$, $\delta z = 0.011$, and $\delta Y = 0.05$, while the relative calibration spread was $\sigma_{68}(g) = 0.014$, $\sigma_{68}(r) = 0.007$, $\sigma_{68}(i) = 0.008$, $\sigma_{68}(z) = 0.013$, and $\sigma_{68}(Y) = 0.015$. These numbers are in good agreement with those quoted above and support the expectation that the relative calibration uncertainty in Table 1 is a conservative estimate.

5. Image Co-addition

Image co-addition allows DES to detect fainter objects and mitigates the impact of residual transient imaging artifacts (e.g., unmasked cosmic rays, satellite streaks, etc.). Combining multiple dithered exposures also positions objects at different points on the focal plane, mitigating systematics associated with the nonuniform response of the instrument.

DESDM produced image co-adds from the weighted average of overlapping single-epoch images. The pixels of the input images were remapped onto a uniform pixel grid using SWarp with the LANCZOS3 kernel (Bertin et al. 2002; Bertin 2010). The remapped pixel grid was defined on co-add *tiles* spanning 0.73 deg \times 0.73 deg and comprising $10^4 \times 10^4$ remapped pixels (a pixel scale of 0.263 pixel⁻¹, comparable to the physical pixel scale of DECam). For each tile, one co-added image was produced for each photometric band.

Before performing image co-addition, several image quality checks were run to identify and blacklist CCD images with severe imaging artifacts. CCD images affected by strong scattered light artifacts were identified by a ray-tracing algorithm using the Yale bright star catalog (Hoffleit & Jaschek 1991), the telescope pointing, and a detailed model of the DECam optics, filter changer, and shutter assemblies. Several exposures have excess noise in one or more of the DECam CCD back planes. These CCD images were identified through visual inspection and through the detection of a large number of spurious catalog objects. In addition, CCD images that were affected by bright meteor trails and airplanes were identified through visual inspection. Less than 1% of CCD images were blacklisted and removed from the co-add process.

⁶² <http://www.stsci.edu/hst/observatory/crds/calspec.html>

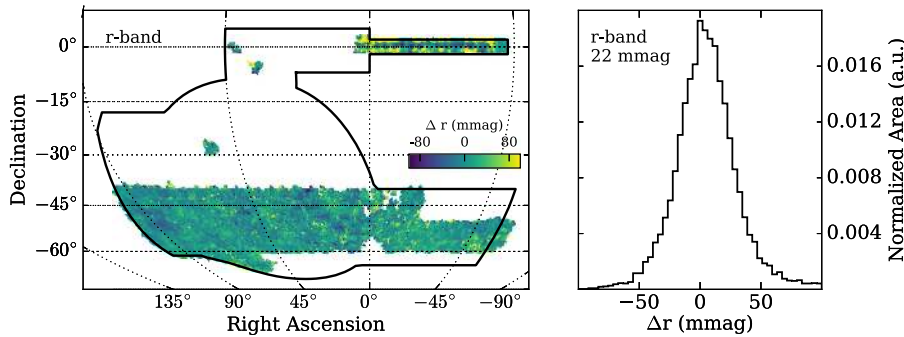


Figure 5. Comparison of stellar magnitudes from the DES Y1A1 GCM and those estimated from APASS/2MASS transformed into the DES filter system (Appendix A.4). The sky plot (left) shows the median magnitude offset for stars binned into $\sim 0.2 \text{ deg}^2$ HEALPIX pixels. The GCM calibrated magnitudes are consistent with the transformed values from APASS/2MASS with a half-width of $\sigma_{68} = 22 \text{ mmag}$ (calculated between the 16th and 84th percentiles). Similar figures for other bands are shown in Appendix A.

When DESDM created co-added images, the PSFs of the individual input images were not homogenized. This decision was motivated by studies of SV data where PSF homogenization was found to produce correlated sky noise, which made it difficult to properly estimate the photometric uncertainties of galaxies. While nonhomogenized PSF co-addition yields better-behaved photometric uncertainties, it can introduce sharp PSF discontinuities on the ~ 0.1 scale that are difficult to model with conventional polynomial approximation techniques (i.e., PSFEX; Bertin 2006). Some of these issues can be addressed by using quantities measured in the Y1A1 FINAL-CUT catalog (Section 6); however, studies that depend sensitively on morphological characterization (i.e., weak lensing analyses) perform their own simultaneous fit of the individual single-epoch images (Section 6.3).⁶³

In addition to the main survey, there are several regions where the DES Y1 imaging is considerably deeper than the nominal three to four tilings. Co-adds have been created in these regions using different numbers of input images to achieve different photometric depths. The Y1A1 GOLD co-add catalog thus contains four different samples:

1. WIDE: The WIDE co-add data sample is built from exposures in the S82 and SPT regions of the Y1 wide-area survey footprint and has a depth of three to four tilings. One of the SN fields, SN-E, resides within the SPT region; however, to maintain uniformity, the WIDE data set only includes images that were taken as part of the DES wide-area survey (the SN-E exposures are included in the other data sets that follow).
2. D04: The D04 sample is constructed by co-adding images in the SN, COSMOS, and VVDS-14h fields with the goal of reaching an effective depth roughly comparable to the WIDE sample. Quantitatively, exposures were selected to give $\sum_j^{\text{exp}} t_{\text{eff},j} T_{\text{exp},j} \simeq 4T_{\text{wide}}$, where $t_{\text{eff},j}$ is the effective exposure timescale factor for exposure j (Section 2), $T_{\text{exp},j}$ is the shutter-open time for exposure j , and T_{wide} is the wide-area exposure time in Y1 (90 s in *griz* and 45 s in *Y*). When selecting exposures for the D04 and D10 samples, we attempted to apply data quality selections based on FWHM and t_{eff} . For the D04 sample, exposures in the *grizY* bands were generally required to pass the

wide-area survey data quality requirements (Section 2) and have $\text{FWHM} < 1''.3$. However, in several cases these requirements were relaxed to better approximate the desired depth. While the D04 sample was designed to mimic the depth of the WIDE survey, the longer exposure times for the auxiliary and SN fields result in a data set that is on average $\sim 0.2 \text{ mag}$ deeper than WIDE. The median $\text{MAG_AUTO } 10\sigma$ limiting magnitude for galaxies (Section 7.1) in the D04 sample is $g = 23.6$, $r = 23.4$, $i = 22.8$, $z = 22.0$, $Y = 20.3$. The D04 data set has been used to train and test photo- z algorithms and object classification (e.g., Hoyle et al. 2017).

3. D10: The D10 sample is constructed in the SN, COSMOS, and VVDS-14h fields by co-adding images to an effective depth of 10 exposures. The 10-exposure depth is intended to mimic the expected main survey depth at the end of DES. Similar to D04, general criteria requiring survey quality, $\text{FWHM} < 1''.3$ in *riz*, and $\text{FWHM} < 1''.4$ in *gY* were applied. The median $\text{MAG_AUTO } 10\sigma$ limiting magnitude for galaxies (Section 7.1) in the D10 sample is $g = 24.2$, $r = 24.0$, $i = 23.5$, $z = 22.7$, $Y = 20.9$.
4. DFULL: The DFULL sample uses all high-quality images in the SN, COSMOS, and VVDS-14h fields. The DFULL co-add applies a requirement of $\text{FWHM} < 1''.3$ in *riz* band and $\text{FWHM} < 1''.4$ in *g* band (no FWHM requirement is placed on *Y* band). Exposures are still required to pass the survey quality cuts, but no restriction is placed on the number of exposures that go into the co-add. The median $\text{MAG_AUTO } 10\sigma$ limiting magnitude for galaxies (Section 7.1) in the DFULL sample is $g = 24.2$, $r = 23.9$, $i = 23.8$, $z = 23.7$, $Y = 21.2$, with $\sim 10\%$ of the area having a limiting magnitude greater than 25 in *griz*.⁶⁴

5.1. Astrometric Accuracy

Astrometric calibration places the DES exposures onto a consistent reference frame with each other and with external catalogs. We used SCAMP (Bertin 2006) to find an astrometric

⁶³ Studies with PSF homogenization are ongoing, and PSF-homogenized co-adds have been used for several DES science analyses using SV data (Hennig et al. 2017; Klein et al. 2017).

⁶⁴ The median depth of the DFULL sample in *g* and *r* band is comparable to that of D10 owing to the fact that few additional exposures passed the survey quality and FWHM requirements outside of the deep SN fields. The *r*-band depth is 0.05 mag shallower in DFULL owing to a slightly larger area with more varied data quality.

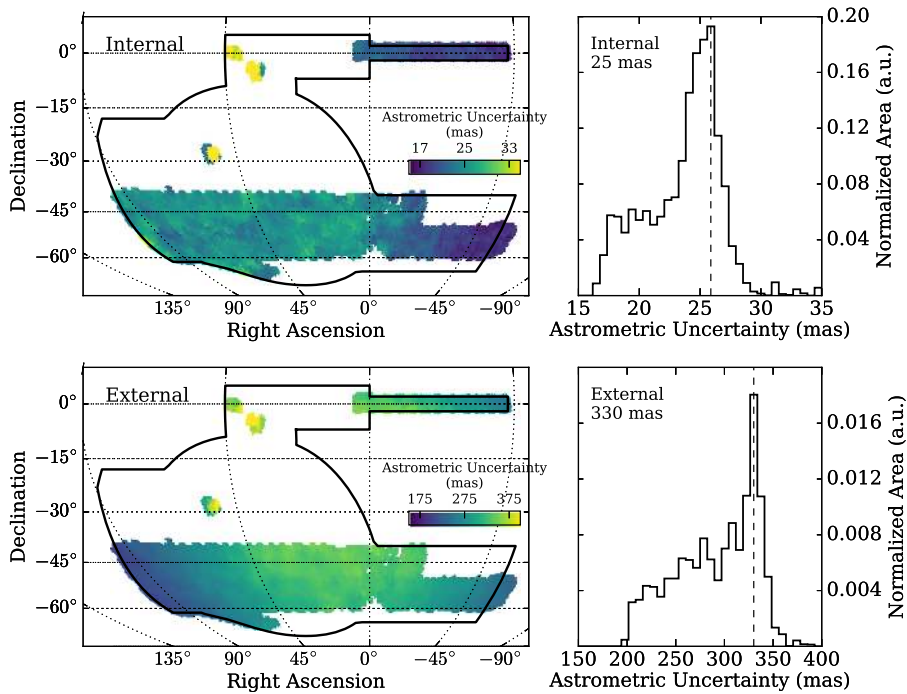


Figure 6. Top: relative internal astrometric error in milliarcseconds derived by comparing the positions of stars in the individual DES exposures that go into the Y1A1 co-adds. Bottom: relative external astrometric error derived by comparing the position of stars in DES and 2MASS (without correcting for proper motion). The color scales represent the astrometric uncertainty in milliarcseconds, while the legends of the right panels report the modes of the distributions. The SN exposure times are significantly longer than the wide-area survey exposure times, leading to a fainter saturation threshold. This reduces the number of nonsaturated bright stars and increases the astrometric uncertainty estimated by this technique (a more accurate estimate of the astrometry in the SN fields can be found in Kessler et al. 2015).

solution including corrections for optical distortion toward the edges of the focal plane. During Y1A1 FINALCUT processing, initial astrometric calibration was performed on individual exposures. Starting with an approximate initial solution provided by the telescope control system, the `SExtractor` windowed image coordinates of bright stars in the DES exposures were extracted and matched against the UCAC-4 stellar catalog (Zacharias et al. 2013).

When building co-add tiles, an additional astrometric refinement process was performed to remap the DES input images against each other and against the 2MASS catalog (Skrutskie et al. 2006). The single-epoch catalogs from all exposures overlapping a tile were input to `SCAMP`, and a simultaneous best fit was obtained treating exposures from each filter as separate instruments. This best-fit astrometric solution was used when combining images. After astrometric refinement, the median internal astrometric precision of the Y1A1 wide-area co-add images is ~ 25 mas (3σ -clipped rms dispersion around the mean for stars with $S/N > 100$). In comparison, the median astrometric precision when compared against 2MASS is 200–350 mas (Figure 6). This difference is dominated by the proper motions of high Galactic latitude stars and uncertainty in the astrometric accuracy of faint 2MASS sources.⁶⁵ This has been confirmed by comparisons between Y3 DES data and *Gaia* DR1 (Gaia Collaboration 2016), where the median astrometric uncertainty is found to be ~ 150 mas (DES Collaboration 2018).⁶⁶

6. Object Catalogs

6.1. Co-add Catalog Creation

Catalogs of unique astrophysical sources were assembled from the co-added images. The goal of the DESDM catalog production was to assemble the most inclusive catalog of sources while maintaining a low contamination fraction. The production of catalog subsamples that are *complete* to a given threshold is left to subsequent science analyses. Source detection, morphological characterization, and multiband photometric flux measurements were performed using `SExtractor` (Bertin & Arnouts 1996; Bertin et al. 2002). Source detection used a CHI-MEAN combination of the co-added images in $r + i + z$ (Szalay et al. 1999; Bertin 2010). The CHI-MEAN detection image was designed to minimize discontinuities between regions with different numbers of exposures (see Appendix B). In contrast, flux and shape measurements were performed on each band individually using `SExtractor` in dual mode (i.e., analyzing the image for an individual band simultaneously with the detection image). The local background was estimated via 16×16 pixel boxes with 3σ clipping of bright pixels and median filtering of the boxes. The image was convolved with a 3×3 pixel structuring element of the form $[[1, 2, 1][2, 4, 2][1, 2, 1]]$. An S/N threshold of 1.5σ per pixel was applied over the convolved image to detect objects. Source localization was derived from the barycenter of the object in the i, z, Y, r, g single-band co-add images (in order). Co-add object positions in world coordinates (J2000 epoch) were computed using the astrometric solution found during image co-addition (Section 5.1).

The depth and PSF of the DES imaging result in overlapping isophotes for objects in crowded regions, e.g., galaxy clusters, star clusters, and dense stellar regions around the LMC.

⁶⁵ http://www.ipac.caltech.edu/2mass/releases/allsky/doc/sec2_2.html

⁶⁶ Bernstein et al. (2017b) show that, using *Gaia* DR1, the astrometric solution for a single DECam exposure can be made accurate to within 3–6 mas.

Table 3
Y1A1 GOLD Catalog Selection

Selection	Description
NITER_MODEL_{GRIZ} > 0 SPREADERR_MODEL_{GRIZ} > 0	Select objects that were observed at least once in each of the g , r , i , z bands. SPREADERR_MODEL = 0 indicates a failure in the photometric fit.

Incomplete deblending of overlapping objects affects the measured shapes and photometric properties of cluster galaxies, which impacts weak lensing and cluster cosmology science. SExtractor attempts to deblend each detected object into subcomponents using a multithresholding algorithm (Bertin & Arnouts 1996). An object is separated into two (or more) new objects if the intensity of the new object is greater than a fraction of the total intensity set by the DEBLEND_MINCONT parameter, while the number of deblending thresholds is set by the DEBLEND_NTHRESH parameter. The Y1A1 processing campaign adopts 0.001 and 32, respectively, for the two parameters. These values were optimized based on SV data to balance completeness and purity for cluster galaxies. More aggressive deblending techniques for the DES data have been explored in Zhang et al. (2015).

SExtractor was used to measure object photometry via several methods (see Sevilla et al. 2011):

1. Fixed aperture fluxes (FLUX_APER) were measured for 12 circular apertures with different radii from $0''.25$ to $9''$.
2. Elliptical aperture fluxes (FLUX_AUTO) were calculated using the second-order moments of each object to derive the elongation and orientation of the best-fit ellipse (Kron 1980). The ellipse scaling factor was derived from the first-order moment of the radial distribution.
3. PSF model fluxes (FLUX_PSF) suitable for point-like sources were fit to the measured PSF shape. As mentioned in Section 5, PSF discontinuities in the Y1A1 co-add images can degrade the quality of the PSF model fluxes.
4. Exponential model fluxes (FLUX_MODEL) suitable for galaxies were fit by convolving a one-component exponential model with a local model of the PSF. These fluxes were fit both individually in each band and by fixing the model shape based on the detection image (FLUX_DETMODEL).

Among the morphological measurements performed by SExtractor, two are designed to separate point-like objects (i.e., stars) from spatially extended sources (i.e., galaxies). The first is the CLASS_STAR variable, which uses a neural network to assess the “stellarity” of an object (Bertin & Arnouts 1996). The second variable, SPREAD_MODEL, is derived from the Fisher’s linear discriminant between a model of the PSF and an extended source model convolved with the PSF (Desai et al. 2012; Bouy et al. 2013; Soumagnac et al. 2015). The application of these variables to star–galaxy separation is detailed in Section 8.1.

As stated previously, catalog quantities were also derived for individual single-epoch exposures that compose the co-added images. Objects detected on the individual exposures were associated with sources in the co-add catalog using a $1''$ matching radius. While shallower, the single-epoch catalogs are important for probing the temporal domain. Additionally, the photometry of the single-epoch catalogs is not subject to the PSF discontinuities present in the co-adds. For this reason, we

calculated a number of photometric and morphological quantities from the average of single-epoch measurements weighted by their associated statistical uncertainties (the names of these quantities are prefixed by “WAVG”). In particular, the weighted-average spread-model quantity (WAVG_SPREAD_MODEL) has been shown to yield better star–galaxy separation (Drlica-Wagner et al. 2015) for stellar objects, and the weighted-average PSF magnitudes (WAVGCALIB_MAG_PSF) have been found to yield more precise stellar photometry than the corresponding co-add quantities. In addition, uncertainties for the WAVG quantities are calculated directly from the variance in the measurements from individual exposures and thus avoid any systematics introduced in the co-addition process.

6.2. Y1A1 GOLD Catalog Selection

We assembled the Y1A1 GOLD object catalog as a high-quality subselection of the objects extracted from the Y1A1 co-add images. When selecting the Y1A1 GOLD catalog, we sought to remove spurious, nonphysical objects while minimally decreasing the statistical power of any scientific investigation (Table 3). Specifically, we required that objects be observed, but not necessarily detected, at least once in each of the g , r , i , and z bands. We also required that all objects have SPREADERR_MODEL > 0 for the g , r , i , and z bands to eliminate objects with unphysical SPREADERR_MODEL values indicative of a failure in the SExtractor photometric fit.⁶⁷ We also identify several classes of objects that are extremely unusual and flag them for exclusion from most cosmological analyses (Table 4). In addition to objects flagged by SExtractor, we specifically identify (1) objects with extremely blue ($\{g - r, r - i, i - z\} < -1$) or extremely red ($\{g - r, r - i, i - z\} > 4$) colors, (2) bright stars that saturate some of the single-epoch inputs to the co-add image, and (3) objects that have a large ($>1''$) offset in the windowed centroid derived from the g and i bands. Finally, we require that objects reside within the Y1A1 GOLD footprint (Section 7.3) and flag any objects that reside in poor-quality or potentially problematic (“bad”) regions (Section 7.4).

6.3. Multiepoch, Multiobject Fitting

The Y1A1 co-added images provide deeper and more sensitive object detection than individual single-epoch images. However, the co-addition process averages across multiple images, resulting in a discontinuous PSF and correlated noise properties. Precision measurements that rely on an accurate PSF determination, such as galaxy shape measurements for cosmic shear, require a joint fit of pixel-level data from multiple single-epoch images.

We used the ngmix⁶⁸ code (Sheldon 2014; Jarvis et al. 2016; Sheldon & Huff 2017) to reanalyze pixel-level data from

⁶⁷ Objects that are not detected in a specific band have a sentinel value of SPREADERR_MODEL = 1.

⁶⁸ <https://github.com/esheldon/ngmix>

Table 4
Y1A1 GOLD Catalog Flags

Flag Bit	Selection	Description
1	$\text{FLAGS}_{\{\text{GRIZ}\}} > 3$	Objects flagged by SExtractor
2	$\{g - r, r - i, i - z\} < -1$ OR $\{g - r, r - i, i - z\} > 4$	Objects with unphysical colors
4	$(\text{NEPOCHS}_G = 0) \text{ AND } (\text{MAGERR_AUTO}_G < 0.05)$ AND $(\text{MAG_MODEL}_I - \text{MAG_AUTO}_I) < -0.4$	Artifacts associated with stars close to the saturation threshold
8	$(\alpha_{J2000,g} - \alpha_{J2000,i} > 1''$ OR $ \delta_{J2000,g} - \delta_{J2000,i} > 1''$) AND $(\text{MAGERR_AUTO}_G < 0.05)$	Objects with large astrometric offsets between bands

multiepoch postage stamps of each object in the Y1A1 GOLD co-add catalog. We used PSFEX (Bertin 2011) to model and interpolate the PSF at the location of each object, and then we generated an image of the PSF using the python package, psfex.⁶⁹ We then used the ngmix code to fit this reconstructed PSF image to a set of three free, independent Gaussians.

We used ngmix in “multiepoch” mode to simultaneously fit a model to all available epochs and bands. In this mode, a model is convolved by the local PSF in each single-epoch image, and a χ^2 sum is calculated over all pixels in a postage stamp. This is repeated for each epoch and band, and a total χ^2 sum is calculated. We then find the parameters of the model that maximize the likelihood.

We took this procedure one step further, performing simultaneous multiepoch, multiband, and multiobject fits, which we call “MOF.” We first identified groups of objects using a friends-of-friends algorithm (e.g., Huchra & Geller 1982; Berlind et al. 2006). We then fit the members of the group using the following procedure:

1. Perform an initial model fit to each object, masking the light from neighbors using the überseg algorithm (Jarvis et al. 2016).
2. Fit the model to each object again, this time subtracting the light from neighbors using the models from the previous fit.
3. Repeat the previous step until all fits converge, or a maximum of 15 iterations was reached.

This fit was performed simultaneously in the g, r, i, z bands using all available imaging epochs and assuming the same spatial model for all bands and epochs. An example of this procedure is shown in Figure 7.

We found that fitting a galaxy model with fully free bulge and disk components was highly unstable, so we adopted the following approach, inspired by the “composite” model used in the SDSS.⁷⁰ We first fit the disk and bulge models separately, represented by an exponential and de Vaucouleurs profile (de Vaucouleurs 1948), respectively. We then determined the linear combination of these models that best fit the data,

$$M_{\text{tot}} = f_{\text{dev}} M_{\text{dev}} + (1 - f_{\text{dev}}) M_{\text{exp}}, \quad (1)$$

where M_{dev} is the bulge model, M_{exp} is the disk model, and f_{dev} represents the fraction of light in the bulge component. This total model is unlikely to be a good fit of the data, and we only

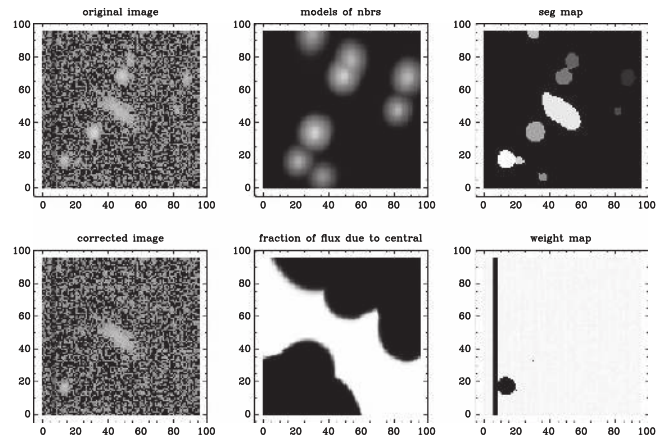


Figure 7. Group of objects fit using the MOF algorithm. In the top row we show the sky-subtracted image (left), the models for neighboring sources (middle), and the SExtractor segmentation map (right). In the bottom row we show the sky-subtracted image after also subtracting the light from neighbors (left), the fraction of light assigned to the central object (middle; 100% in white and 0% in black), and the weight map (right). Note that a bad column and a flagged object are identified in the weight map. The masked object was not fit, and thus its light was not subtracted.

use it as a starting point for a more refined model. We formed a new model that has the best f_{dev} determined as above, as well as the same ratio of scale lengths for the bulge and disk components. This new model has free parameters for the center, ellipticity, overall scale, and fluxes. A common center, scale, and ellipticity were used for all bands, but the flux for each band was left free.

For computational efficiency, each component of this model was approximated by a sum of Gaussians (Hogg & Lang 2013). This choice made convolution with the triple-Gaussian PSF model very fast. A fast approximation for the exponential function was also used to speed up computations (Sheldon 2014).

We imposed uninformative priors on all parameters except for the ellipticity and the fraction of light present in the bulge, f_{dev} . For both of these parameters, we applied priors based on fits to deep COSMOS imaging data, provided as postage stamps with the GalSim project.⁷¹ We defined convergence to be when the flux from subsequent fits to objects did not change more than a part in a thousand, and structural parameters such as scale and ellipticity did not change by more than one part in a million. For incorporation into the Y1A1 GOLD catalog, we converted MOF fluxes to magnitudes and applied the SLR adjustment discussed in Section 4.3.

⁶⁹ <https://github.com/esheldon/psfex>

⁷⁰ <http://www.sdss.org/dr12/algorithms/magnitudes/#cmodel>

⁷¹ <https://github.com/GalSim-developers/GalSim>

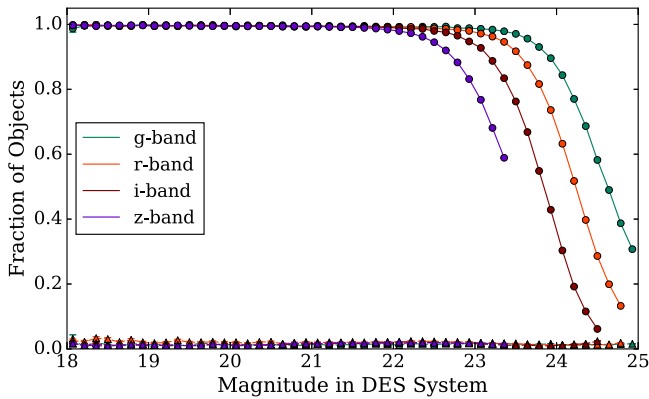


Figure 8. Completeness (solid circles) and contamination (dashed triangles) of the Y1A1 GOLD co-add object catalog determined by comparison to the CFHTLenS W4 field. Object matching was performed within a $1''$ radius, and CFHTLenS magnitudes were transformed to the DES system using the equations in Appendix A.4. Statistics were calculated for the subset of objects that were unmasked in both surveys and have been truncated at the 5σ limiting magnitude of CFHTLenS (Erben et al. 2013).

6.4. Catalog Completeness

We assessed the completeness and purity of the Y1A1 GOLD catalog by comparing it against data from the Canada–France–Hawaii Lensing Survey (CFHTLenS) W4 field (Hildebrandt et al. 2012; Erben et al. 2013), which overlap the S82 region of Y1A1 GOLD. The DES data in this overlap region have a typical 10σ limiting magnitude of $g \sim 23.1$, $r \sim 23.0$, $i \sim 22.5$, $z \sim 21.8$ (Section 7.1). This is comparable to the median for Y1A1 GOLD in i and z bands and ~ 0.2 mag shallower than the median in g and r bands (Table 1). In this region, CFHTLenS is $\gtrsim 1$ mag deeper than the Y1A1 GOLD catalog, making it a good test for object detection completeness. We transformed the magnitude of CFHTLenS objects into the DES system (Appendix A.4) and removed objects residing in masked regions of either survey. We associated objects between the two catalogs based on a spatial coincidence of $1''$ and required a matching magnitude within 2 mag. We then calculated the detection completeness as the fraction of CFHTLenS objects that are matched to Y1A1 GOLD objects as a function of the CFHTLenS magnitude transformed into the DES system. The contamination of the Y1A1 GOLD catalog is assessed as the fraction of Y1A1 GOLD objects that are unmatched to CFHTLenS objects as a function of magnitude. We find that the 95% completeness limit of the Y1A1 GOLD catalog is $g = 23.6$, $r = 23.4$, $i = 22.9$, and $z = 22.4$ (Figure 8). We find that for magnitudes brighter than these limits, the contamination of the Y1A1 GOLD catalog is $\lesssim 2\%$. The Y1A1 GOLD catalog is $>99\%$ complete in all four bands for magnitudes brighter than 21.5. This completeness estimate does not account for objects that are blended in both CFHTLenS and DES, which is estimated to be $\sim 1\%$ of objects at DES depth. We also note that Y1A1 object detection was performed on a combined $r + i + z$ detection image and no S/N threshold was applied to the measurements in individual bands when calculating completeness.

7. Ancillary Maps

Several ancillary maps were produced to characterize the coverage, sensitivity, observing conditions, and potentially problematic regions of Y1A1 GOLD as a function of sky

position. Generating ancillary maps for Y1A1 GOLD was a multistep process: we created a vectorized representation of the survey coverage and limiting magnitude using `mangle` (Hamilton & Tegmark 2004; Swanson et al. 2008), we rasterized the `mangle` maps with `HEALPix` for ease of use, we estimated observing conditions over the survey footprint, and we subselected a nominal high-quality footprint. Finally, we flagged sky regions where the true survey performance deviates from that estimated by the ancillary data products (i.e., the regions around bright stars, astrometric failures, etc.). Each of these steps is described in more detail below.

7.1. Maps of Survey Coverage and Depth

Quantifying survey coverage and limiting magnitude as a function of sky position are essential for statistically rigorous cosmological analyses. To accurately track characteristics of the DES survey at the sub-CCD level, DESDM produces `mangle` masks (Hamilton & Tegmark 2004; Swanson et al. 2008) as part of the Y1A1 COADD pipeline. These masks are an accurate representation of the coverage, sensitivity, and overlap of DECam exposures including dead CCDs, gaps between CCDs, masked regions around bright stars, and bright streaks from Earth-orbiting satellites.

During co-add production, `mangle` masks were created at the level of co-add tiles (Figure 9). The steps are as follows:

1. Polygons were created using the four corners of each input CCD image and assigned a weight equal to the median value of pixels in the CCD weight plane.
2. Satellite streaks were represented by polygons, and the area of these polygons was removed from the single-epoch CCD polygon.
3. Polygons were trimmed to fit the tile boundaries.
4. Polygons were subdivided into disjoint regions with the `balkanize` command. Following the weighted-average scheme chosen for image co-addition, the total weight of a balkanized polygon is the sum of the weights of the individual polygons.
5. Regions around bright stars and bleed trails are removed from the `mangle` mask. While the precise location of these artifacts is image dependent, it is computationally simpler to mask the stacked map with the largest shape covering a bright star or bleed trail rather than removing these regions from each single-epoch polygon.
6. The `mangle` co-add weight map was converted into a 10σ limiting magnitude map for a $2''$ diameter aperture:

$$m_{\text{lim}} = m_{\text{ZP}} - 2.5 \log \left(10 \sqrt{\pi \frac{(D/2)^2}{\omega_{\text{pix}}^2} \frac{1}{w_{\text{tot}}}} \right), \quad (2)$$

where $m_{\text{ZP}} = 30$ is the tile zero-point, $D = 2''$, $\omega_{\text{pix}} = 0''.263$ is the pixel size, and w_{tot} is the total weight of the polygon. This definition of the magnitude limit corresponds to the `MAG_APER_4` quantity measured by `SExtractor`.

While the vectorized `mangle` masks are a very accurate representation of the DES survey coverage, they are computationally unwieldy for many scientific analyses. To increase the speed and ease with which survey coverage and limiting magnitude can be accessed, we generate anti-aliased `HEALPix` maps of these quantities (Figure 9). Pixelized maps of the survey coverage fraction were created at a resolution of

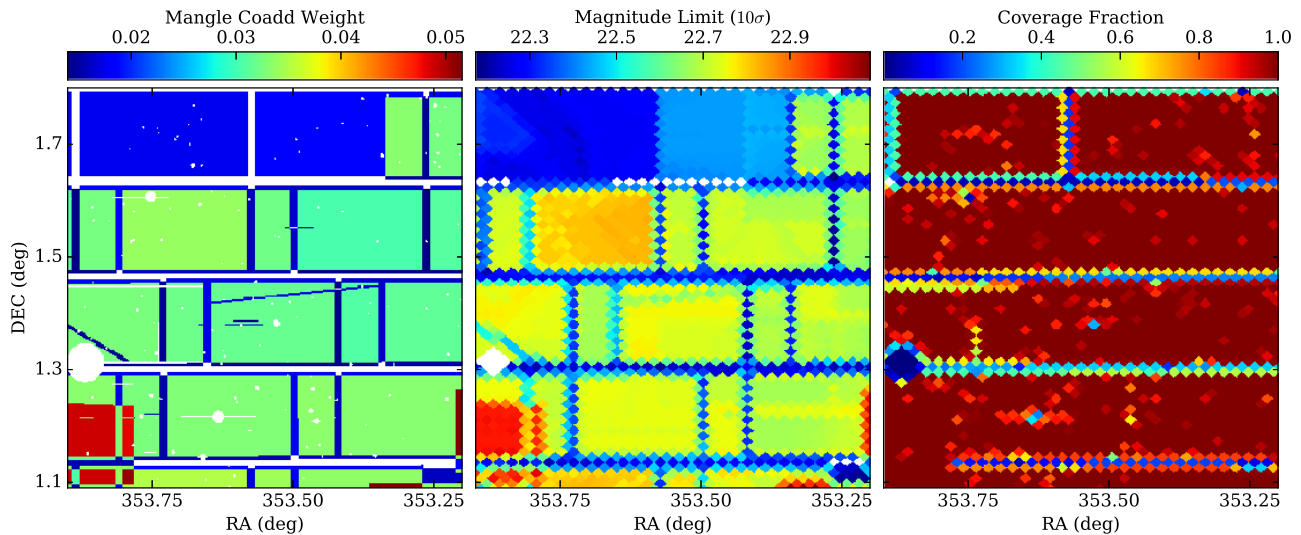


Figure 9. Coverage and depth maps for a single Y1A1 co-add tile. Left: vectorized `mangle` weight map for an r -band co-add tile. Satellite trails, star masks, and chip gaps are stored at full resolution. Middle: pixelized 10σ limiting magnitude map for galaxies using HEALPix at $n_{\text{side}} = 4096$. Right: pixelized map of the coverage fraction at HEALPix $n_{\text{side}} = 4096$. This tile is located on the border of the Y1A1 footprint and has been chosen for illustrative purposes owing to its variable depth and incomplete coverage.

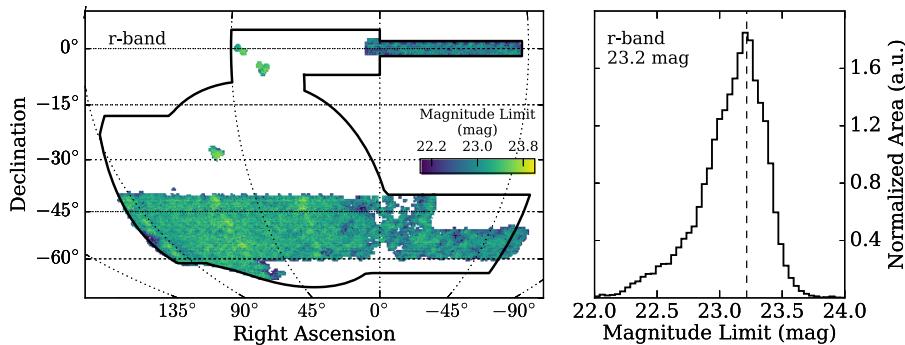


Figure 10. Sky map and normalized histogram for the r -band 10σ limiting magnitude (`MAG_AUTO`) derived in HEALPix pixels over the Y1A1 GOLD footprint. The mode of the limiting magnitude distribution is shown in the right panel. The derivation of the limiting magnitude is described in Section 7.1. Similar figures for other bands are shown in Appendix C.

$n_{\text{side}} = 4096$ (area = 0.73 arcmin^2) by calculating the fraction of higher-resolution subpixels ($n_{\text{side}} = 32768$, area = 0.01 arcmin^2) that were contained within the `mangle` mask. Similarly, maps of the survey limiting magnitude were generated at $n_{\text{side}} = 4096$ by calculating the mean limiting magnitude for subpixels ($n_{\text{side}} = 32768$). When calculating the limiting magnitude, subpixels that were not covered by the survey were excluded from the calculation, while subpixels that have been masked (i.e., bright stars, bleed trails, etc.) were assumed to have the limiting magnitude of their parent polygon. The HEALPix resolution of $n_{\text{side}} = 4096$ was chosen as a compromise between computational accuracy and ease of use. This resolution was found to have a negligible effect on the correlation function of simulated galaxies on scales larger than 0.5 when combined with the survey coverage fraction maps.

We followed the prescription of Rykoff et al. (2015) to convert the `mangle` coverage and depth maps into 10σ limiting magnitude maps for galaxy photometry. We selected galaxies using the `MODEST_CLASS` star–galaxy classifier (Section 8.1) and trained a random forest model to predict the 10σ limiting magnitude as a function of observing conditions. The input vector for the random forest included

the PSF FWHM, sky brightness, airmass, and exposure time for each band being fit (Section 7.2). The training was performed on coarse HEALPix pixels ($n_{\text{side}} = 1024$) that contained more than 100 galaxies. Once trained, the model was applied to the pixels at the full mask resolution of $n_{\text{side}} = 4096$. We derived magnitude limits for both co-add `AUTO` magnitudes and the multiepoch composite model magnitudes derived by the MOF (Section 6.3). We applied the SLR calibration adjustment (Section 4.3) to the resulting depth maps to correct for interstellar extinction and zero-point nonuniformity. The median 10σ limiting magnitudes for `MAG_AUTO` are $g = 23.4^{+0.14}_{-0.40}$, $r = 23.2^{+0.15}_{-0.37}$, $i = 22.5^{+0.14}_{-0.34}$, $z = 21.8^{+0.12}_{-0.37}$, $Y = 20.1^{+0.18}_{-0.33}$, where the uncertainties represent the 16th and 84th percentiles of the distribution. In comparison, the median 10σ limiting magnitudes for the MOF `CM_MAG` magnitudes are $g = 23.7^{+0.07}_{-0.40}$, $r = 23.5^{+0.16}_{-0.29}$, $i = 22.9^{+0.14}_{-0.30}$, and $z = 22.2^{+0.14}_{-0.32}$. We find that the depth estimates are accurate at the level of 6%–7%, but that 3%–4% of this measured uncertainty is due to “pixelization noise” resulting from averaging over a range of depths when fitting the model on coarse pixels. An example of the resulting depth maps for r band can be found in Figure 10, and figures for the other bands can be found in Appendix C.

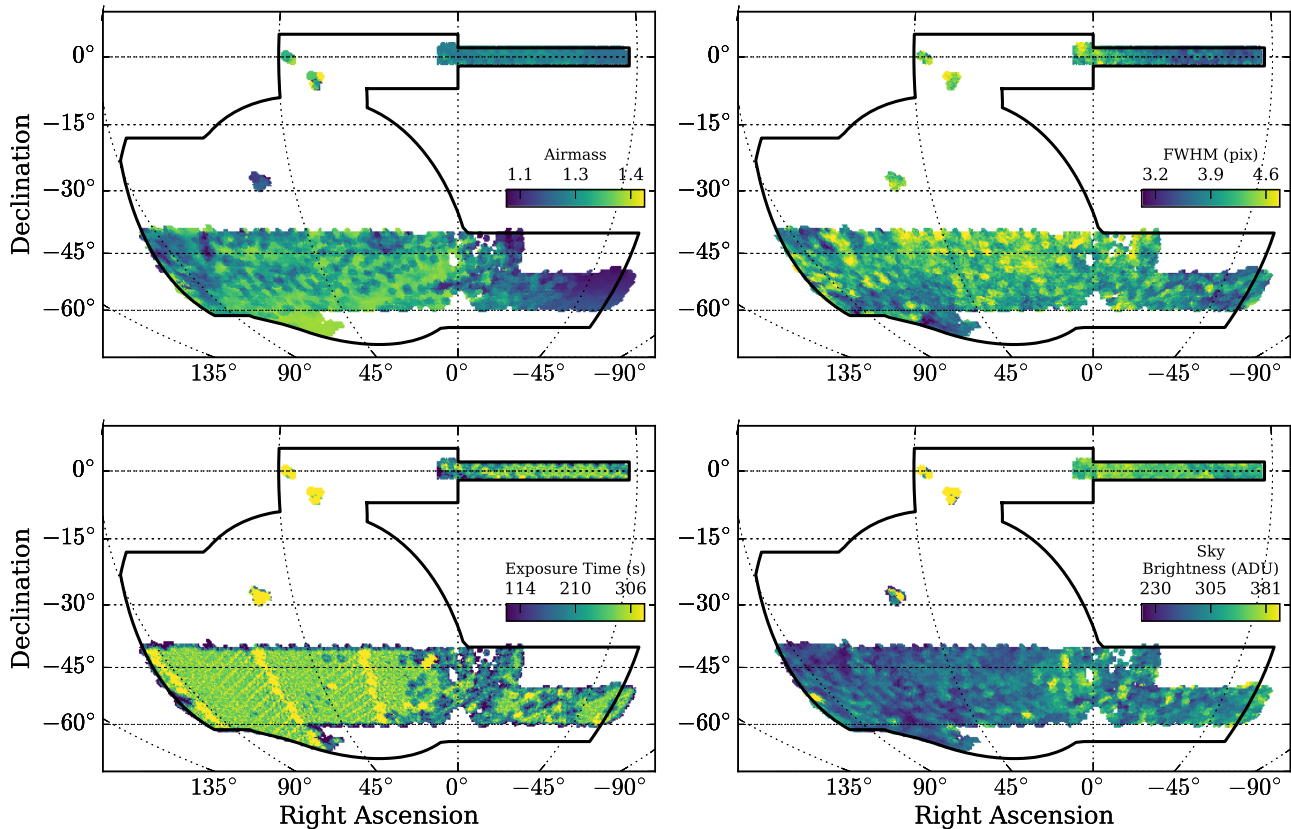


Figure 11. Survey characteristics of the Y1A1 GOLD data set estimated from the inverse variance weighted stack of single-epoch images in r band at each position on the sky. Panels correspond to mean airmass (top left), PSF FWHM in pixels (top right), exposure time in seconds (bottom left), and sky brightness in ADU (bottom right).

7.2. Maps of Survey Characteristics

Variations in observing conditions can be a significant source of systematic uncertainty in cosmological analyses. In a wide-area optical survey such as DES, variable observing conditions can imprint spurious spatial correlations, noise, and depth fluctuations on the object catalogs that are used for galaxy clustering and cosmic shear analyses. By identifying and characterizing these systematic effects, it becomes possible to quantify and minimize their impact on scientific results. We followed the procedure developed by Leistedt et al. (2016) to construct survey characteristic and coverage fraction maps for the Y1A1 GOLD data set using QuickSip.⁷² Since the nonlinear transfer function between the stack of images at any position on the sky and the final galaxy catalog is largely unknown, we created maps of many different survey observables. For each band, we created maps of both weighted- and unweighted-average quantities of each image. The main quantities expected to be used for null tests in cosmological analyses with the Y1A1 GOLD catalog are the total exposure time, the mean PSF FWHM, the mean airmass, and the sky background. The inverse variance weighted averages of these quantities are shown in Figure 11. Further modeling of the survey transfer function is important for DES cosmology analyses, and several approaches have already been developed (e.g., Chang et al. 2015; Suchyta et al. 2016).

7.3. Footprint Map

The nominal footprint for the Y1A1 GOLD catalog is defined using an $n_{\text{side}} = 4096$ HEALPix map. For a pixel to be included in the Y1A1 GOLD footprint, it must meet the following criteria simultaneously in the g , r , i , z bands:

1. A mangle coverage fraction ≥ 0.5 , implying that at least half of the pixel area has been observed or is unmasked according to mangle (Section 7.1).
2. A coverage fraction of ≥ 0.5 from the survey characteristics maps (Section 7.2).
3. A minimum total exposure time of ≥ 90 s (Section 7.2).
4. A valid solution from the SLR calibration adjustment (Section 4.3).

These selection criteria reduce the total co-added area of Y1A1 covered in any band, 1927 deg^2 , to a nominal WIDE+D04 Y1A1 GOLD footprint in g , r , i , z of 1786 deg^2 . Simultaneously applying the same criteria to the Y band (with a minimum exposure time of 45 s) results in a g , r , i , z , Y footprint of 1773 deg^2 . These numbers were calculated by summing the coverage fraction of pixels in the footprint.

7.4. Bad Region Mask

Masks were developed to remove regions where survey artifacts make it difficult to control systematic uncertainties when doing cosmological analyses. Since not all science topics require the same masks (e.g., studies of galaxy evolution may not want to mask nearby galaxies), the various masks are

⁷² <https://github.com/ixkael/QuickSip>

Table 5
Y1A1 GOLD Bad Region Mask

Flag Bit	Area (deg ²)	Description
1	30.1	High density of astrometric discrepancies
2	119.5	2MASS moderate star regions ($8 < J < 12$)
4	5.4	RC3 large galaxy region ($10 < B < 16$)
8	38.6	2MASS bright star regions ($5 < J < 8$)
16	95.8	Region near the LMC
32	18.4	Yale bright star regions ($-2 < V < 5.6$)
64	1.3	High density of unphysical colors
128	...	Unused bit
256	0.7	Milky Way globular clusters
512	7.2	Poor COADD PSF modeling

Note. Masked regions for the Y1A1 GOLD WIDE+D04 footprint. The masked area is calculated using the coverage fraction of the pixels that are removed from the footprint by each mask. The criteria defining each mask can be found in Section 7.4.

collected into a bitmap defined in Table 5. Removing area associated with any of these masks results in a WIDE+D04 footprint area of 1506 deg² in g, r, i, z and 1496 deg² in g, r, i, z, Y .

7.4.1. Catalog Artifacts

1. *Unphysical colors (bit = 64)*: This mask is designed to remove imaging artifacts that were not masked before creating co-adds. In particular, this mask removes regions where there are significant reflected light artifacts (both specular and diffuse) from bright stars, unmasked orbital satellite trails, and co-add saturation artifacts. This mask is pixelized at $n_{\text{side}} = 2048$, and pixels with ≥ 8 objects possessing unphysical colors are masked (see Table 4). The threshold for flagging bad pixels was set by visual inspection of the co-add tiles. The resulting masked area is 1.3 deg².
2. *Astrometric discrepancies (bit = 1)*: We flag regions that have a high concentration of galaxies with large astrometric offsets between filters. We select galaxies with $i < 22$, $\text{MAGERR_AUTO_G} < 0.2$, and windowed positions in g and i band differ by more than $1''$. This criterion has been found to select objects in regions of strongly variable background (e.g., the wings of bright stars, regions of poor sky subtraction, regions with scattered light, etc.). The resulting masked area in this case is 30.1 deg².
3. *PSF model failures (bit = 512)*: There are several regions where PSF modeling failed owing to varying depth and a discontinuous PSF. Co-add tiles possessing poor PSF models are identified as having a large number of stars where the co-add PSF magnitude differs from the weighted-average single-epoch PSF magnitude by more than 0.2 mag. We flag HEALPIX pixels ($n_{\text{side}} = 512$) possessing > 20 stars with large discrepancies in PSF magnitudes. The total region masked in this way is 7.2 deg².

7.4.2. Bright Stars

Regions around saturated stars were masked at the pixel level as part of the image processing pipeline described in

Section 3. However, catalog-level investigation revealed a residual increase in the number density of objects surrounding the brightest stars. To avoid contamination from spurious objects in the halos of bright stars, we designed radial masks based on the brightness of the contaminating stars and the number density of surrounding objects. These masks were developed for two bright star catalogs as described below.

1. *Yale bright star regions (bit = 32)*: Masked regions were determined from the positions and magnitudes of stars in the Yale Bright Star Catalog (Hoffleit & Jaschek 1991). The masking radius was determined from the V -band magnitude of each star, following the equation

$$R = 0.86835 \text{ deg} - (0.1439 \text{ deg}) \times V. \quad (3)$$

Minimum and maximum masking radii were imposed at 0.1 deg and 0.4 deg, respectively. The resulting masked area is 18.4 deg².

2. *2MASS bright stars (bit = 2, 8)*: We mask regions around bright stars from the 2MASS catalog (Skrutskie et al. 2006) within a radius of

$$R = 0.09 \text{ deg} - (0.0073 \text{ deg}) \times J, \quad (4)$$

assuming a minimum and maximum masking radius of 0.01 deg and 0.05 deg, respectively. Many of the bright stars in 2MASS overlap with the faintest stars in the Yale Bright Star Catalog, and we find a comparable masking radius (albeit derived using different bands). Because the fainter 2MASS stars may not be problematic for all science applications, we split the 2MASS star mask into stars with $5 < J < 8$ and stars with $8 < J < 12$. The masked areas are 38.6 deg² and 119.5 deg², respectively.

7.4.3. Large Foreground Objects

1. *The Large Magellanic Cloud (bit = 16)*: The center of the Large Magellanic Cloud (LMC) is located ~ 5 deg from the southwest edge of the DES footprint, and the stellar population of the LMC presents a number of challenges for extragalactic science. The high density of stars decreases the purity of galaxy samples, while the 2MASS star masks described in Section 7.4.2 lead to a complex and heavily masked area. The stellar locus of the LMC differs from that of the Milky Way, making it difficult to apply the SLR calibration adjustment described in Section 4. For these reasons, we masked a region around the LMC with a boundary defined as $60^\circ < \alpha_{2000} < 100^\circ$ and $-70^\circ < \delta_{2000} < -58^\circ$. The LMC mask removed 95.8 deg².
2. *Bright galaxies (bit = 4)*: The Third Reference Catalog of Bright Galaxies (RC3; Corwin et al. 1994) contains galaxies subtending $\gtrsim 1'$. Since galaxy size is highly correlated with magnitude, we continue to use a magnitude-dependent masking formulation similar to that applied to bright stars. We masked a circular region around RC3 galaxies with $10 < B < 16$ with a magnitude-dependent selection:

$$R = 0.269 \text{ deg} - (0.0166 \text{ deg}) \times B. \quad (5)$$

We imposed minimum and maximum masking radii such that $0.03 \text{ deg} < R < 0.1 \text{ deg}$. The bright galaxy mask removes 5.4 deg².

Table 6
Y1A1 GOLD MODEST_CLASS Star–Galaxy Classification

Class	Selection	Description
0	$\text{SPREAD_MODEL_I} + (5/3) \times \text{SPREADERR_MODEL_I} < -0.002$	Unphysical PSF fit (likely stars)
1	$\text{SPREAD_MODEL_I} + (5/3) \times \text{SPREADERR_MODEL_I} > 0.005$ AND NOT ($ \text{WAVG_SPREAD_MODEL_I} < 0.002$ AND $\text{MAG_AUTO_I} < 21.5$)	High-confidence galaxies
2	$ \text{SPREAD_MODEL_I} + (5/3) \times \text{SPREADERR_MODEL_I} < 0.002$	High-confidence stars
3	$0.002 < \text{SPREAD_MODEL_I} + (5/3) \times \text{SPREADERR_MODEL_I} < 0.005$	Ambiguous classification

Note. The high-purity and high-completeness galaxy samples are defined as $\text{MODEST_CLASS} = 1$ and $\text{MODEST_CLASS} \in \{1, 3\}$, respectively. Similarly, the high-purity and high-completeness stellar samples are defined as $\text{MODEST_CLASS} = 2$ and $\text{MODEST_CLASS} \in \{0, 2, 3\}$, respectively.

3. *Globular clusters (bit = 256):* The high stellar density of Milky Way globular clusters makes them difficult regions for cosmology analyses. We identified three globular clusters, NGC 1261, NGC 1851, and NGC 7089, and masked circular regions with radius 1.5 times the angular size reported by Sinnott (1988). This resulted in a total masked area of 0.7 deg².

8. Value-added Quantities

The astrometric, photometric, and morphological parameters derived for each object are supplemented with additional information important for astrophysical and cosmological analyses. These “value-added quantities” are built from the calibrated co-add object catalog and provide additional information on an object-by-object basis. The two primary value-added quantities provided with Y1A1 GOLD are (1) a simple star–galaxy classifier and (2) a set of photo-*z* estimates.

8.1. Star–Galaxy Separation

As part of the Y1A1 GOLD catalog, we produced a “MODEST_CLASS” object classification with the primary goal of selecting high-quality galaxy samples. MODEST_CLASS is based on the *i*-band co-add quantity SPREAD_MODEL_I and its associated error, SPREADERR_MODEL_I. SPREAD_MODEL is a morphological variable defined as a normalized linear discriminant between the best-fit local PSF model and a slightly more extended model composed of a circular exponential disk convolved with the PSF (Desai et al. 2012; Soumagnac et al. 2015). The *i* band was chosen as the reference band for object classification owing to its depth and superior PSF. Image-level simulations of the DES data support the conclusion that *i* band yields the best overall performance for object classification, and this result was verified using deep *HST* imaging on the COSMOS field.

We used space-based imaging of COSMOS (Leauthaud et al. 2007) and GOODS-S (Giavalisco et al. 2004) along with spectroscopic observations from VVDS (Le Fèvre et al. 2005) that overlapped the Y1A1 GOLD footprint as a truth sample for developing MODEST_CLASS. We defined star and galaxy samples optimized for “high completeness” and “high purity” by applying thresholds on the combination of SPREAD_MODEL_I and SPREADERR_MODEL_I.⁷³ The object classification scheme is defined in Table 6 and shown graphically in Figure 12.

Following Drlica-Wagner et al. (2015), we validated the performance of the MODEST_CLASS star–galaxy classifier on

data from CFHTLenS (Hildebrandt et al. 2012; Erben et al. 2013). We matched CFHTLenS catalog objects to the Y1A1 GOLD data (Section 6) and selected high-quality samples of stars and galaxies using the `class_star` and `fitclass` measurements by CFHTLenS (Heymans et al. 2012). Specifically, our CFHTLenS stellar selection was (`fitclass = 1`) OR (`class_star > 0.98`) and our galaxy selection was (`fitclass = 0`) OR (`class_star < 0.2`). Note that ~7% of matched CFHTLenS objects are unclassified according to this prescription, and these objects are not used for assessing the performance of MODEST_CLASS.

We define the “efficiency” of a galaxy sample as the number of true galaxies that are also classified as galaxies divided by the total number of true galaxies in the sample (i.e., the true positive rate). Conversely, the “contamination” of a galaxy sample is defined as the number of galaxies that are misclassified divided by the total number of objects classified as galaxies (i.e., the false discovery rate). Similar definitions apply to the stellar selections, and the performance of the MODEST_CLASS galaxy and star selections are shown in Figure 13. We find that a high-purity galaxy selection has an efficiency $\gtrsim 98\%$ and a contamination rate $\lesssim 3\%$ for $i < 22$. In contrast, the high-completeness stellar selection has an efficiency of $\gtrsim 86\%$ with a contamination of $\lesssim 6\%$ for $i < 22$. We estimate similar performance for MODEST_CLASS through a comparison against the DEEP2-3 field in the first public data release of Hyper Suprime Camera (Aihara et al. 2017).

The MODEST_CLASS selection provides an initial baseline for object classification and is found to be sufficient for characterizing the distributions of stars and galaxies in Y1A1 GOLD (Figures 14 and 15). Multivariate machine-learning techniques and template-fitting algorithms have the potential to provide much better object classification (e.g., Fadelly et al. 2012; Soumagnac et al. 2015, etc.). Several advanced object classification techniques are currently being explored within DES and will be detailed in future publications (I. Sevilla et al. 2017, in preparation). We emphasize that MODEST_CLASS has been optimized for *galaxy* selection. Several alternative selections have been suggested for more complete samples of stars (e.g., Bechtol et al. 2015; Drlica-Wagner et al. 2015).

8.2. Photometric Redshift Estimation

In this section we briefly summarize the approach to photo-*z* estimation and validation for DES Y1 science analyses. While photo-*z* estimates were provided as part of the initial Y1A1 GOLD data set, it was realized that individual cosmology analyses benefit from photo-*z* estimation and validation customized to their distinct science samples. Therefore, we present a general overview of the photo-*z* estimation and validation procedures, and we refer the reader to upcoming

⁷³ The high-completeness and high-purity samples differ in the classification assigned to ambiguous objects.

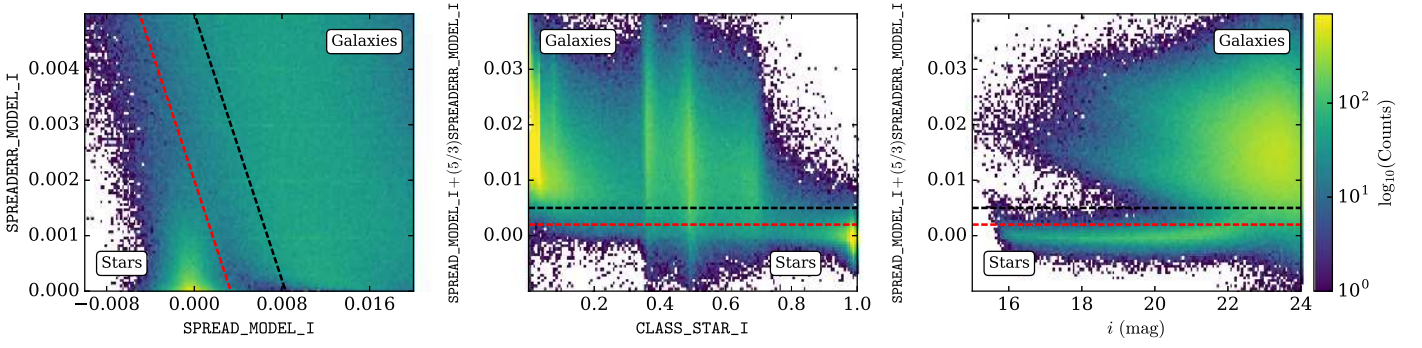


Figure 12. MODEST_CLASS star–galaxy selection for objects in a $\sim 13 \text{ deg}^2$ region centered on $\alpha_{2000}, \delta_{2000} = (51^\circ, -45^\circ)$. The left panel shows the distribution of `spread_model_i` and its error. The middle panel compares the distribution of the SExtractor neural-network classifier, `class_star`, with the MODEST_CLASS selection criteria. The right panel shows a tight stellar locus in the *i*-band magnitude compared against the MODEST_CLASS criteria. In all panels the black (red) lines correspond to the pure (complete) galaxy selection threshold applied on MODEST_CLASS.

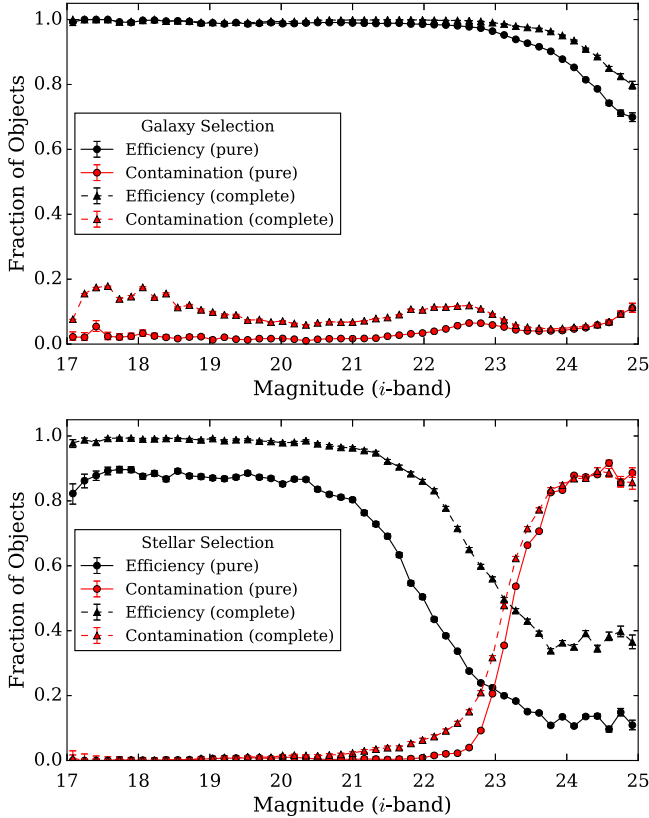


Figure 13. Performance of the MODEST_CLASS star–galaxy classifier based on a comparison to deeper imaging from CFHTLenS. Top: measured efficiency and contamination for high-purity (solid) and high-completeness (dashed) galaxy samples. Bottom: efficiency and contamination for high-purity (solid) and high-completeness (dashed) stellar samples (note that MODEST_CLASS is not optimized for stellar selection).

publications dedicated to photo-*z* estimation for distinct DES analyses (e.g., Cawthon et al. 2017; Davis et al. 2017; Gatti et al. 2018; Hoyle et al. 2017).

Photo-*z* estimates were generated with two distinct algorithms: the machine-learning code DNF (De Vicente et al. 2016), and a modified version of the template code BPZ (Benéz 2000; Hoyle et al. 2017). These two codes are representative of common machine-learning and template-fitting photo-*z* estimation techniques. Both algorithms utilized spectroscopic data for training, and a detailed discussion of

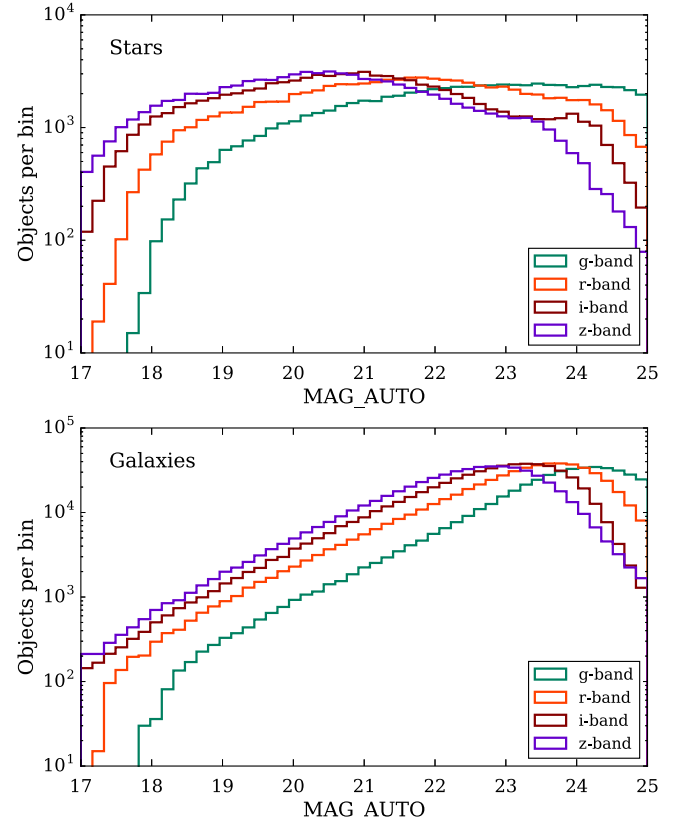


Figure 14. Number counts of objects passing the MODEST_CLASS pure star selection (top) and complete galaxy selection (bottom) as a function of `MAG_AUTO` magnitude in the *g*, *r*, *i*, and *z* bands. The impact of galaxy contamination can be seen in the stellar number counts at magnitudes fainter than $i \gtrsim 23$.

the spectroscopic sample can be found in Gschwend et al. (2017).

For many cosmological analyses, we are interested in accurately characterizing the statistical distribution of galaxies in tomographic bins of redshift and less interested in predicting the redshift of any individual galaxy. Thus, we applied two independent techniques targeted at validating the statistical properties of our predicted photo-*z* distributions (Davis et al. 2017; Hoyle et al. 2017).

1. We performed a direct validation of the color–redshift relationship by matching galaxies from DES science

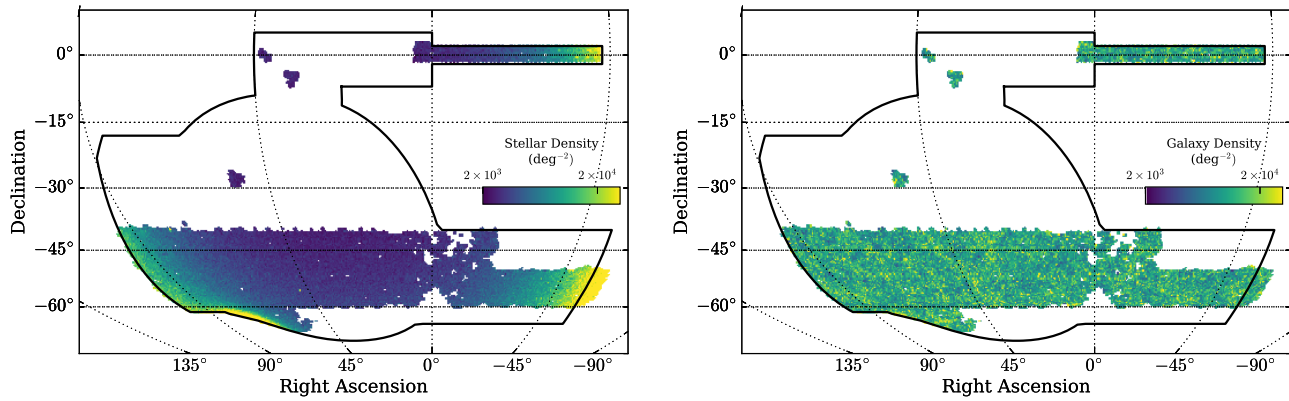


Figure 15. Density of objects with $i < 22$ passing the high-completeness star (left) and high-purity galaxy (right) MODEST_CLASS selections. The linear color scales represent the density of catalog objects and are the same for both panels. The density of objects has been corrected for the coverage fraction of each pixel (Section 7.1).

samples to galaxies with multiband photometry obtained within the COSMOS field (Laigle et al. 2016). This choice of validation data mitigated the impact of redshift or galaxy-type dependent selection biases, which can affect spectroscopic surveys (e.g., Bonnett et al. 2016; W. Hartley et al. 2018, in preparation). However, the 30-band photo- z estimates from COSMOS have a larger intrinsic uncertainty than spectroscopically determined redshifts. In addition, validating performance on a $\sim 2 \text{ deg}^2$ field leads to large uncertainty due to cosmic variance, which was estimated using the Buzzard suite of Λ CDM simulations (J. DeRose et al. 2017, in preparation; Sánchez et al. 2017; R. H. Wechsler et al. 2018, in preparation).

2. A second, independent indirect validation technique relies on the clustering-redshift technique (Newman 2008; Ménard et al. 2013; Schmidt et al. 2013). We selected a luminous red galaxy sample (redMaGiC; Rozo et al. 2016), which has well-determined photo- z estimates, as a reference and divided this sample into redshift bins of width $\Delta z = 0.02$. We then divided the full sample of DES objects into tomographic redshift bins based on predicted photo- z and cross-correlated the data in each tomographic bin with each of the more finely binned redMaGiC reference samples. We measured the excess angular cross-correlation signal, which is proportional to the redshift distribution. We calibrated a constant redshift offset in each tomographic bin between the photo- z predictions and the clustering signal. We estimated the errors arising from the evolution of galaxy–dark matter halo bias and discrepancies in the shape of the clustering redshift distribution by repeating the same analysis using the Buzzard simulations (Cawthon et al. 2017; Gatti et al. 2018).

Both validation techniques possess associated uncertainties. The direct validation technique has comparable uncertainties from sample variance (since COSMOS covers a $\sim 2 \text{ deg}^2$ region of the sky) and systematic uncertainty in matching the morphological and color–magnitude error distribution of the galaxy sample. In contrast, we find that the dominant systematic uncertainties for the indirect validation technique come from the clustering bias evolution of the binned source galaxy samples and incorrectness in the shape of the photo- z

distribution. In addition, we are unable to perform indirect clustering validation for tomographic bins with $z \gtrsim 1$ owing to limited redMaGiC reference data at these redshifts.

The most important photo- z performance metric for cosmic shear analyses is the bias of the estimated mean of a redshift distribution in a tomographic bin with respect to the unknown true mean redshift in that bin (Bonnett et al. 2016). We characterized the photo- z accuracy from the photo- z bias distribution, defined as the difference between the average measured photometric redshift and the average true redshift distribution, $\Delta z = \langle z_{\text{true}} \rangle - \langle z_{\text{phot}} \rangle$. Since the true redshift distribution is unknown, we employed the direct and indirect validation techniques described above to estimate $\langle z_{\text{true}} \rangle$ and Δz in four tomographic bins with $0.2 < z < 1.3$. We find that both techniques yield $|\Delta z| \lesssim 0.02$ with an uncertainty of comparable magnitude when applied to the BPZ estimates for the primary subsample of the Y1A1 GOLD catalog used for cosmic shear analyses (Hoyle et al. 2017; Zuntz et al. 2017).

We present several other results from the validation of the BPZ template code optimized over the redshift range $0.2 < z < 1.3$ for the primary Y1 weak-lensing shear catalog (Zuntz et al. 2017). In Figure 16, we show the $n(z)$ distribution for the weak-lensing shear catalog derived from Y1A1 GOLD. The $n(z)$ distribution is found to be in good agreement with the $n(z)$ predicted from COSMOS when cosmic variance and other associated systematic uncertainties are accounted for (Hoyle et al. 2017). We also show a comparison between the redshift estimate from a random sampling of the 30-band COSMOS $P(z)$ (Laigle et al. 2016) and the median photo- z derived from DES Y1 using BPZ. Structure along the line of sight is visible in the higher-resolution COSMOS redshifts but is not resolved by DES. For the full weak-lensing subsample, the normalized median absolute deviation (NMAD) of the quantity $(z_{\text{DES}} - z_{\text{COSMOS}})/(1 + z_{\text{COSMOS}})$ is 0.08–0.09, depending on the point estimate used to determine the DES BPZ redshift. When restricted to $i \leq 22$, the photo- z NMAD decreases to 0.06–0.07. Due to the strict selection requirements of the weak-lensing subsample, the NMAD for the full Y1A1 GOLD galaxy sample is slightly larger (~ 0.12). The photometric redshift accuracy for forthcoming DES Y1 cosmology analyses will be documented in more detail in Hoyle et al. (2017), Cawthon et al. (2017), Davis et al. (2017), and Gatti et al. (2018).

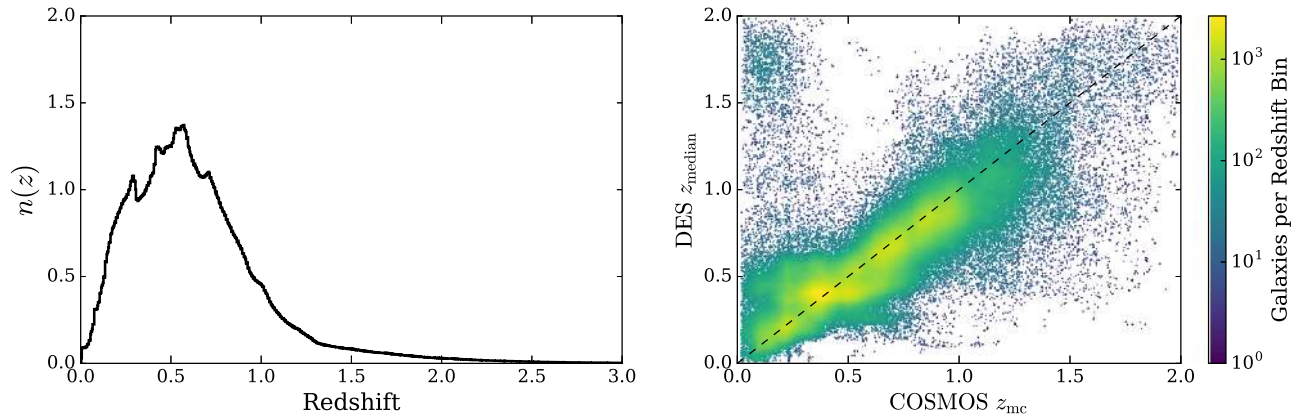


Figure 16. Performance of the BPZ photo- z estimates for the weak-lensing subsample of Y1A1 GOLD. Left: $n(z)$ distribution generated from a random Monte Carlo sampling (i.e., Wittman 2009) of the BPZ $P(z)$ distribution for each galaxy. Right: comparison between a random sampling of the 30-band COSMOS $P(z)$ (Laigle et al. 2016) and the median of the DES BPZ $P(z)$ for each galaxy in the overlapping sample. Points are colored by the local density of comparison galaxies in 0.04×0.04 redshift bins.

9. Conclusion

During its first year, DES imaged $\sim 2000 \text{ deg}^2$ of the southern sky in each of the g , r , i , z , and Y photometric filters. These data have been processed, calibrated, co-added, cataloged, and characterized to form the DES Y1A1 GOLD cosmology data sample, which covers $\sim 1800 \text{ deg}^2$ with a depth of three to four tilings per band and a photometric calibration accuracy of $\lesssim 2\%$. The photometric calibration uniformity of the Y1A1 GOLD catalog was validated and adjusted via an SLR technique, which also corrects for the effects of interstellar extinction on the calibrated magnitudes of objects. The development of Y1A1 GOLD was driven by the goal of producing a maximal sample of Y1 data while minimizing the impact of systematic features. Several ancillary maps characterizing the DES survey and its performance were produced as part of Y1A1 GOLD. In addition, a simple star-galaxy classifier and several photometric redshift estimates were also produced as necessary precursors to many DES science analyses. The Y1A1 GOLD data set is intended to be used as the nominal starting point for cosmological analyses with the DES Y1 data.

The next DES co-added data set will consist of exposures from the first three seasons of DES and will increase both the survey coverage ($\sim 5000 \text{ deg}^2$) and depth (five to six tilings per band). Improvements to the Blanco telescope infrastructure, data processing algorithms, and photometric calibration are expected to yield higher-quality data. In addition, many of the improvements developed for Y1A1 GOLD have been integrated into the core DESDM processing pipeline (e.g., Morganson et al. 2018) and into automated tools for science catalog creation (e.g., A. Fausti Neto et al. 2017, in preparation). However, we anticipate that future data sets will still require the construction and validation of a high-quality data sample to serve as the basis for cosmological analyses. On the longer term, we expect that a similar procedure for assembling and validating cosmology data samples will be necessary for future surveys, such as LSST. We hope that the production of Y1A1 GOLD will help serve as a road map for assembling cosmology-ready data samples for future large photometric surveys.

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The UKIDSS project is defined in Lawrence et al. (2007). UKIDSS uses the UKIRT Wide Field Camera (WFCAM; Casali et al. 2007). The photometric system is described in Hewett et al. (2006), and the calibration is described in Hodgkin et al. (2009). The pipeline processing and science archive are described in M. J. Irwin et al. (2009, in preparation) and Hambly et al. (2008).

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Facility: Blanco (DECam).

Software: SExtractor (Bertin & Arnouts 1996), PSFEx (Bertin 2011), SCAMP (Bertin 2006), SWarp (Bertin et al. 2002; Bertin 2010), mangle (Hamilton & Tegmark 2004; Swanson et al. 2008), HEALPix (Górski et al. 2005),⁷⁴ astropy (Astropy Collaboration 2013), matplotlib (Hunter 2007), numpy (Van Der Walt et al. 2011), scipy (Jones et al. 2001), healpy,⁷⁵ fitsio,⁷⁶ ngmix (Sheldon 2014).⁷⁷

Appendix A Photometric Calibration

In this appendix, we provide more details on the photometric calibration of Y1A1, including nightly calibration (Appendix A.1), global calibration (Appendix A.2), and an SLR adjustment (Appendix A.3). We note that the nightly and global calibration steps followed on the procedure of Tucker et al. (2007) and were performed on the single-epoch catalog data *before* co-addition. In contrast, the SLR adjustment was performed on the weighted-average magnitudes of multiple single-epoch catalogs and is applied directly to the co-added object catalogs. A collection of transformation equations between DES and several other surveys is provided in Appendix A.4.

A.1. Nightly Photometric Calibration

The first step in DES Y1 photometric calibration used observations of standard-star fields to derive a set of calibration coefficients for each photometric night. A subset of the standard-star fields listed in Table 7 were observed at different airmasses at the beginning and end of each DES night or half night (Section 2). The DES nightly standard-star fields are predominantly located in the equatorial fields of SDSS Data Release 9 (DR9; Ahn et al. 2012), with the addition of several fields from the Southern $u'g'r'i'z'$ Standard Network (Smith et al. 2017).⁷⁸ Devoted observations of these standard-star fields were supplemented by DES survey observations that overlapped the standard-star fields. For Y band, we used stars from the equatorial fields of the UKIRT Infrared Deep Sky Survey Data Release 6 (UKIDSS DR6; Lawrence et al. 2007) matched against SDSS stars. All nightly standard stars were transformed to an initial DES AB photometric system via matching to objects in SDSS and UKIDSS (Appendix A.4). The spatial distribution of the DES standard-star fields is shown in Figure 17.

Due to its provenance, primarily from SDSS DR9, we refer to the set of DES nightly standards as *secondary* standards. The *fundamental* standard for SDSS was the F subdwarf star BD +17°4708, which was used for calibrating the set of SDSS *primary* standards (Smith et al. 2002). These primary standards were in turn used (indirectly) in the ubercalibration of SDSS (Padmanabhan et al. 2008). Thus, the DES secondary standards tie the absolute flux calibration of DES to the SDSS primary standards, to BD +17°4708, and ultimately to the AB magnitude system.

Nightly observations of the secondary standards were used to fit a set of photometric equations. These photometric equations, which are based on those used by SDSS (Tucker

⁷⁴ <http://healpix.sourceforge.net>

⁷⁵ <https://github.com/healpy/healpy>

⁷⁶ <https://github.com/esheldon/fitsio>

⁷⁷ <https://github.com/esheldon/ngmix>

⁷⁸ <http://www-star.fnal.gov>

Table 7
DES Nightly Standard-star Fields

Field Name	R.A. J2000	Decl. J2000	Exposure Time (s)				
			<i>g</i>	<i>r</i>	<i>i</i>	<i>z</i>	<i>Y</i>
<i>Preferred Fields^a</i>							
SDSS J2140-0000	21:40:00	+00:00:00	15	15	15	15	20
SDSS J2300-0000	23:00:00	+00:00:00	15	15	15	15	20
SDSS J0000-0000	00:00:00	+00:00:00	15	15	15	15	20
SDSS J0100-0000	01:00:00	+00:00:00	15	15	15	15	20
SDSS J0200-0000	02:00:00	+00:00:00	15	15	15	15	20
SDSS J0320-0000	03:20:00	+00:00:00	15	15	15	15	20
SDSS J0843-0000	08:43:00	+00:00:00	15	15	15	15	20
SDSS J0933-0005	09:33:00	-00:05:00	15	15	15	15	20
SDSS J0958-0010	09:58:00	-00:10:00	15	15	15	15	20
SDSS J1048-0000	10:48:00	+00:00:00	15	15	15	15	20
SDSS J1227-0000	12:27:00	+00:00:00	15	15	15	15	20
SDSS J1442-0005	14:42:00	-00:05:00	15	15	15	15	20
C26202/ <i>HST</i>	03:32:30	-27:46:05	15	15	15	15	20
MaxVis	06:30:00	-58:45:00	15	15	15	15	20
<i>Supplemental Fields^b</i>							
SA E1-A	01:24:50	-44:33:40	3	3	3	3	5
SA E2-A	04:03:00	-44:41:45	3	3	3	3	5
SA E3-A	06:42:54	-45:05:06	3	3	3	3	5
SA E4-A	09:23:44	-45:21:02	3	3	3	3	5
SA E5-A	12:04:11	-45:24:03	3	3	3	3	5
SA E6-A	14:45:33	-45:15:34	3	3	3	3	5
SA E8-A	20:07:22	-44:37:01	3	3	3	3	5
SA E9-A	22:45:37	-44:22:47	3	3	3	3	5

Notes.

^a The preferred fields (with the exception of C26202/*HST* and MaxVis) have photometric standard stars covering the entire DECam focal plane. This permits photometric zero-points to be determined for every CCD using a single exposure.

^b The supplemental fields have photometric standard stars covering a $10' \times 10'$ region and are typically used for expanding the range of airmasses when no suitable primary field is observable. These particular supplemental fields come from the Southern *ugriz* Standard Stars project (Smith et al. 2017).

et al. 2006), have the form

$$g_0 = -2.5 \log_{10}(F_g) - a_g - b_g \times ((g - r)_0 - (g - r)_{\text{fid}}) - k_g \times X, \quad (6)$$

$$r_0 = -2.5 \log_{10}(F_r) - a_r - b_r \times ((g - r)_0 - (g - r)_{\text{fid}}) - k_r \times X, \quad (7)$$

$$i_0 = -2.5 \log_{10}(F_i) - a_i - b_i \times ((i - z)_0 - (i - z)_{\text{fid}}) - k_i \times X, \quad (8)$$

$$z_0 = -2.5 \log_{10}(F_z) - a_z - b_z \times ((i - z)_0 - (i - z)_{\text{fid}}) - k_z \times X, \quad (9)$$

$$Y_0 = -2.5 \log_{10}(F_Y) - a_Y - b_Y \times ((z - Y)_0 - (z - Y)_{\text{fid}}) - k_Y \times X, \quad (10)$$

where λ_0 ($\lambda = g, r, i, z, Y$) is the calibrated standard-star magnitude in the DES system, F_λ is the observed PSF flux (counts s^{-1}), a is the photometric zero-point for the night, b is the instrumental color term coefficient, $(g - r)_0$, $(i - z)_0$, and

$(z - Y)_0$ are the calibrated standard-star colors, $(g - r)_{\text{fid}}$, $(i - z)_{\text{fid}}$, and $(z - Y)_{\text{fid}}$ are fiducial reference colors (chosen so that the effects of b are relatively small for a star of typical color within the DES footprint), k is the first-order extinction coefficient, and X is the airmass of the observation. The values for the fiducial colors are $(g - r)_{\text{fid}} = 0.53$, $(i - z)_{\text{fid}} = 0.09$, and $(z - Y)_{\text{fid}} = 0.05$.

Separate a and b coefficients were determined for each functioning science CCD, while a single value of k was assumed for the full focal plane. The nightly values of a track the overall throughput of the DECam instrument at the location of each CCD, and variations in a mostly track the gradual accumulation of dust on the Blanco primary mirror. The nightly values of b track variations in the *shape* of the total filter response curve at the location of each CCD (including atmospheric transmission). Under photometric conditions, the value of k should not vary across the focal plane, and a single value of k was fit for the full focal plane. Variations in the nightly values of k track the relative throughput of the atmosphere at CTIO. The median values of a and b are shown for each science CCD in Figure 18, and the nightly variations of a , b , and k are shown in Figure 19. The site-average values for the a , b , and k coefficients are tabulated in Table 8.

We note that, despite the use of star flats and pupil corrections, there are still minor variations in the zero-points across the focal plane in Figure 18. These variations can be attributed to two main sources: (1) the DES star flat procedure is subject to small flat/planar gradients across the focal plane, with the understanding that the photometric calibration procedure will remove such gradients; and (2) CCD-to-CCD variations in quantum efficiency have not been fully accounted for in the Y1A1 image processing, and this is also reflected in the smaller-scale between-CCD variations in the zero-points. The DES Y3 processing has largely corrected for variations in quantum efficiency, while Burke et al. (2017) show that gradients in the star flats can be successfully removed by the photometric calibration.

The code used to perform the nightly fits is called the Photometric Standards Module (PSM).⁷⁹ We note that PSM not only fits Equations (6)–(10) but also performs an automated initial culling of nonphotometric data using the outputs of RASICAM (Lewis et al. 2010; Reil et al. 2014). PSM also culls dome-occulted exposures (identified by a strong gradient in a across the focal plane) and performs iterative sigma-clipping to achieve a good solution for a night. For the Y1 data set, we also culled nights with rms fit residuals > 0.025 mag or an atypical ($\sim 2\sigma$ – 3σ outlier) fit value for the first-order extinction. The typical relative calibration scatter from the PSM solution is ~ 0.02 mag rms. This scatter includes the contribution from stellar shot noise in the standard-star observations.

A.2. Global Calibration

In addition to deriving the nightly calibration coefficients for each photometric night in DES Y1, we would like to calibrate exposures taken under cloudy conditions and exposures where the nightly solution failed (e.g., due to contrails). We would also like to improve on the $\sim 2\%$ rms relative calibration uncertainty achieved by the PSM solution. To achieve both of

⁷⁹ <https://github.com/DarkEnergySurvey/PSM>

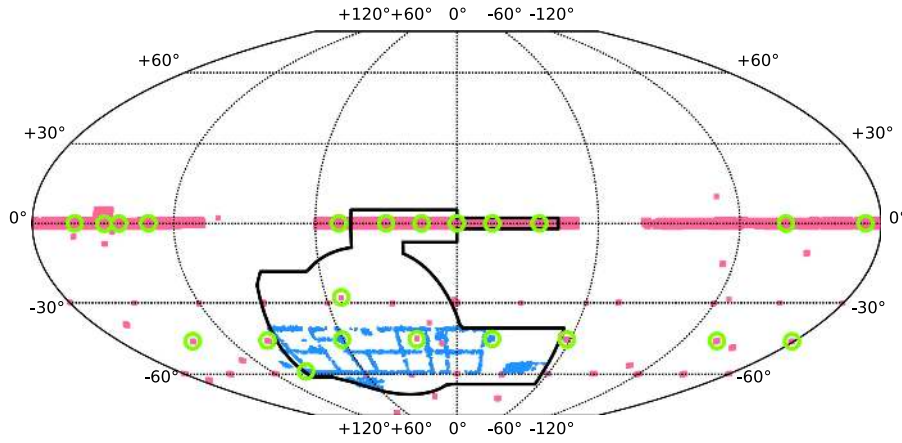


Figure 17. Standard stars used for the photometric calibration of DES Y1. Nightly standard-star fields (Table 7) are marked with green circles, while other secondary standards from SDSS and Smith et al. (2017) are shown in pink. The grid work of tertiary standards is shown in blue.

these goals, we applied a global calibration to simultaneously calibrate all overlapping CCD images in the Y1 data set. In addition to calibrating images that lacked a PSM solution, the global calibration can achieve a relative calibration between overlapping images at the level of 0.003 mag (0.3%) rms, even if the exposures were taken under cloudy conditions.

The global calibration was implemented as a Global Calibrations Module (GCM).⁸⁰ The GCM generalizes the procedure of Glazebrook et al. (1994) by replacing overlapping image “frames” with arbitrarily shaped overlapping catalog data sets (these are still conventionally referred to as “images”). The procedure is summarized briefly as follows.

1. For each filter, consider n data sets for which $(1, \dots, m)$ are uncalibrated and $(m + 1, \dots, n)$ are calibrated. In most cases these data sets represent object catalogs from individual exposures or CCD images. However, the calibrated data set consists of standard stars spanning the entire Y1A1 footprint (Figure 17).
2. Compile a list of all unique pairs of observations of a common star on two data sets.
3. For a given pair of images, i, j , let

$$\Delta_{ij} = \text{median}_{\text{pairs}}(m_i - m_j), \quad (11)$$

where m_i is the magnitude of a star in image i , m_j is the magnitude of the same star in image j , and the median is calculated over matched pairs of stars. Note that $\Delta_{ij} = -\Delta_{ji}$.

4. Let ZP_i be a floating zero-point that can be applied to the data set from image i to produce calibrated magnitudes. For images that are already calibrated ($i > m$), we fix $ZP_i = 0$.
5. Let θ_{ij} define a function that selects overlapping image pairs. We define $\theta_{ij} = 1$ if $i = j$ or if i and j overlap; otherwise, $\theta_{ij} = 0$.
6. To find calibrated zero-points for each image, we minimize the sum of squares,

$$S = \sum_{i,j} \theta_{ij} (\Delta_{ij} + ZP_i - ZP_j)^2. \quad (12)$$

7. We derive a calibrated magnitude for each object detected on image i (where $i < m$) by adding ZP_i to the raw instrumental magnitudes.

In Figure 20, we show a simple example of the GCM algorithm on two disconnected groups of three overlapping data sets (i.e., images). In each group, one of the overlapping images has been previously calibrated and serves as the reference against which the other images in its grouping are calibrated. To be calibrated, an uncalibrated image needs either to overlap a calibrated image (e.g., the left group in Figure 20) or to have an unbroken path of overlapping images to a calibrated image (e.g., image 3 in the right panel of Figure 20). In the right panel of Figure 20 we show the matrix equation that minimizes Equation (12) for this particular set of images (Glazebrook et al. 1994). Note that, via this matrix equation, the zero-points for the two calibrated images (images 5 and 6) have been fixed to a value of zero ($1 \times ZP_5 = 0$ and $1 \times ZP_6 = 0$), since no offset is applied to these previously calibrated images.

Following the prescription of Glazebrook et al. (1994), we estimate the rms magnitude residual for each CCD image, i , from overlap with other CCD images, j , as

$$\text{rms}_i = \sqrt{\frac{\sum_j \theta_{ij} (\Delta_{ij} + ZP_i - ZP_j)^2}{\sum_j \theta_{ij}}}. \quad (13)$$

The rms distribution over all CCD images is a measure of the internal (reproducibility) errors on small scales (the scales of overlapping CCD images) and is a measure of the precision of the overall GCM solution.

In principle, the GCM method is very precise but carries the caveat that any small systematic gradients in the flat-fielding of individual images can cause low-amplitude gradients over large scales. We used the set of secondary standard stars and a sparse grid work of tertiary standard stars (Figure 17) as an “anchor” to keep the GCM fit from drifting owing to any small systematic gradients in the Y1A1 FINAL CUT exposures. We note that the sparse grid work of tertiary standards in the SPT region was extracted from stars in photometric exposures, calibrated by the nightly PSM results. From the full set of PSM-calibrated exposures in the SPT region, we selected a sparse grid work of thin (one-degree-wide) “struts” of constant R.A. and constant decl. This morphology was chosen so that the tertiaries would anchor the GCM solution on large scales (>10 – 15 deg), but on smaller scales the calibration would be dominated by the GCM solution for overlapping uncalibrated exposures.

⁸⁰ <https://github.com/DarkEnergySurvey/GCM>

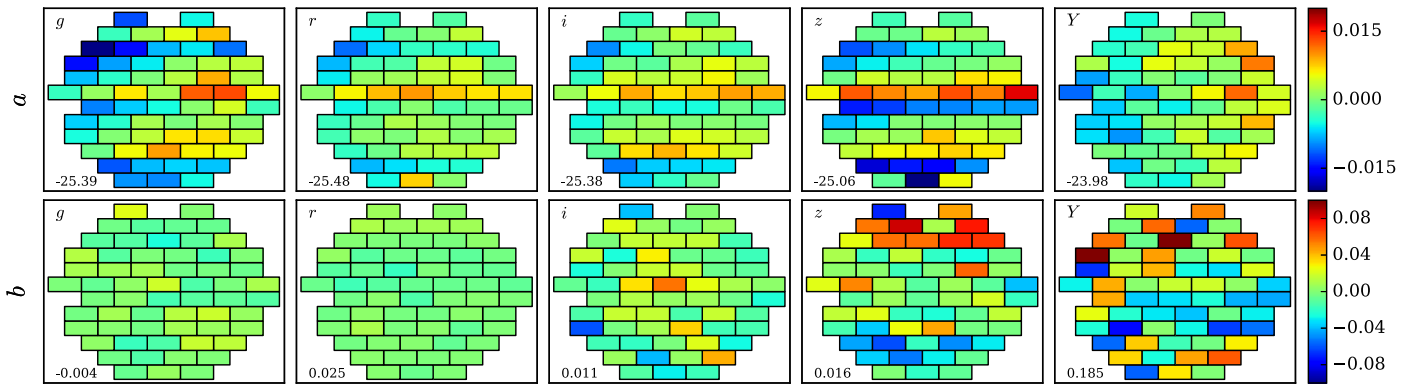


Figure 18. Median fit values from DES Y1 for the a and b coefficients as a function of position on the DECam focal plane. The color scale represents the offset in the median fit value for each CCD with respect to the median for the focal plane as listed in the bottom left of each panel. The spatial variations in photometric zero-points across the DECam focal plane are typically less than 0.02–0.03 mag for Y1A1.

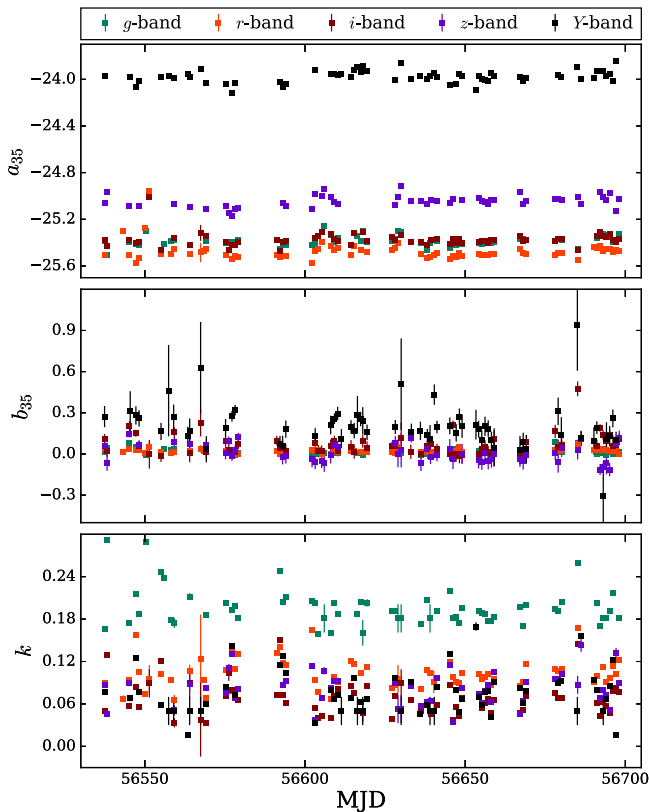


Figure 19. Nightly fit values for the a coefficient of CCD35 (top), the b coefficient of CCD35 (middle), and the k coefficient of the DECam focal plane (bottom) for nights in DES Y1. For the a and b coefficients, the trends in CCD35 are found to be representative of the trends in the other CCDs.

The GCM algorithm relies on having at least one calibrated image or data set to anchor each isolated image group. To identify isolated groups of exposures, we employed a group-finding algorithm developed for studies of galaxy clusters and large-scale structure (Huchra & Geller 1982). For Y1A1 FINALCUT, there were several disconnected image groups—in particular, the SPT region, the S82 region, the four SN fields (SN-E was treated as an isolated group even though it overlaps the SPT area), the COSMOS field, and the VVDS-14h field (see Figure 2). We therefore ran GCM separately on each of these eight regions. The S82, COSMOS, and VVDS-14h fields overlap with the equatorial region of SDSS and were anchored by the secondary standards (mostly derived from SDSS). SPT

Table 8
DES Y1 Average PSM Fit Values

Coeff.	Band	Median	Mean	σ	Mean Error	Number
a	g	-25.396	-25.389	0.044	0.001	3717
	r	-25.501	-25.485	0.079	0.001	4072
	i	-25.386	-25.380	0.059	0.001	3806
	z	-25.056	-25.056	0.049	0.001	3080
	Y	-23.976	-23.979	0.051	0.001	3627
b	g	-0.004	-0.004	0.017	<0.001	3717
	r	0.024	0.024	0.014	<0.001	4072
	i	0.012	0.012	0.057	0.001	3806
	z	0.020	0.020	0.071	0.001	3080
	Y	0.187	0.187	0.123	0.002	3627
k	g	0.191	0.196	0.027	0.003	63
	r	0.099	0.102	0.021	0.003	69
	i	0.065	0.068	0.024	0.003	64
	z	0.083	0.081	0.023	0.003	52
	Y	0.070	0.075	0.031	0.004	61

Note. Statistics were calculated for nights in Y1 with a good PSM fit. The a and b values were calculated individually for each CCD, while the k values were calculated for the full focal plane.

was anchored by the aforementioned grid work of tertiaries, supplemented with individual fields from DES SV. The SN fields also used individual standard-star fields from SV for their calibrators. As with the grid work of tertiary standards, the individual SV fields had been previously calibrated using nightly results from the PSM code (Wyatt et al. 2014).

The S82 region and the smaller individual fields (COSMOS, VVDS-14h, and the SN fields) were each calibrated with a single pass of the GCM. This run treated the catalog from each individual CCD image as the unit to be calibrated and yielded zero-point offsets for each CCD directly. Due to its large area, the SPT region was calibrated from multiple iterations of the GCM. The first pass treated the full catalog from each exposure (59 or 60 functioning science CCDs) as the unit to be calibrated. This could be done because, due to the star flat procedure, all the CCDs on a given exposure have very nearly the same zero-point (at least for exposures taken under photometric conditions). In this pass, small (2%–3%) variations in the relative zero-point across the focal plane were temporarily removed using the median a coefficients for each CCD (Figure 18). In this manner, each exposure was temporarily flat-fielded across the focal plane to reduce

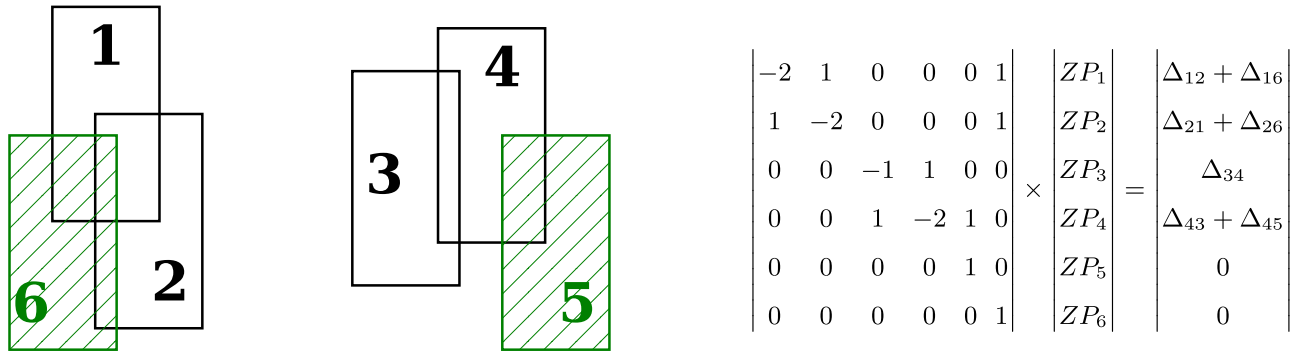


Figure 20. Schematic of the GCM algorithm based on Figure 1 of Glazebrook et al. (1994). Left: the stars in images 5 and 6 have been previously calibrated, while the stars in the other images are uncalibrated. The algorithm minimizes the zero-point offsets from all the overlapping images. Images that have a connected path via overlapping images to a reference image can be calibrated to that reference image. Right: corresponding matrix equation for this set of images.

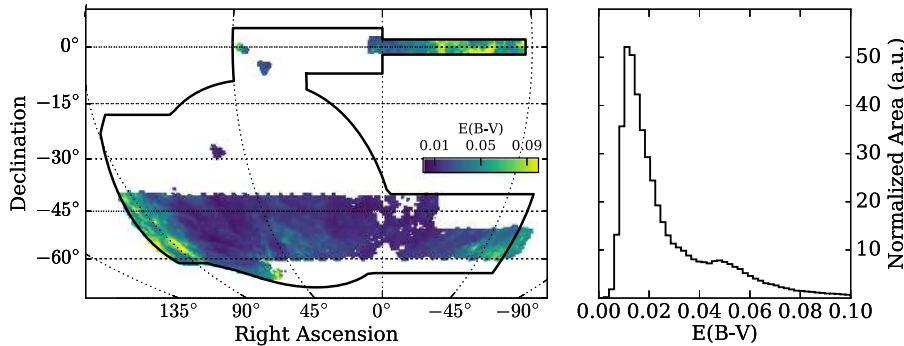


Figure 21. Interstellar extinction, $E(B - V)$, over the Y1A1 GOLD footprint taken from Schlegel et al. (1998). The DES footprint was explicitly chosen to occupy a low-extinction region at high Galactic latitude.

exposure-scale photometric gradients. For the first pass, only exposures that were classified as having been observed under photometric conditions—as determined by RASICAM—were allowed in the GCM fit. The first pass yielded a set of zero-point offsets—one per exposure—for all the (apparently) photometric exposures in the SPT region. The second run of GCM was essentially identical to the first, but it removed outlier exposures—ones with particularly “noisy” or discrepant zero-points. For both the first and second runs, the sparse grid work of tertiaries and the handful of individual calibrated SV fields (Figure 17) were used as the calibrated data set for the Glazebrook et al. (1994) algorithm. Again, this yielded a set of zero-point offsets—one per exposure—for all the photometric exposures in the SPT region. These individual CCD zero-point offsets were applied to all the CCD images in the set of photometric exposures included in the second-pass run of GCM, creating a set of “quaternary” standard stars covering nearly all of the SPT region. In the third and final run of the GCM for the SPT region, the catalog from each individual CCD image was treated—as in the case of GCM runs for S82, COSMOS, VVDS-14h, and the SN fields—as the unit to be calibrated. Furthermore, *all* CCD images from the SPT region—those from photometric exposures and those from nonphotometric exposures—were included in the GCM fit. For this third pass of the GCM, the newly created quaternary standard stars were used as the calibrated data set. This third pass of the GCM for the SPT region yielded a set of zero-point offsets for each CCD image, which was used to calibrate the Y1A1 GOLD single-epoch CCD images in advance of the image co-addition process.

A.3. Photometric Calibration Adjustment

To correct for residual color nonuniformity in the photometric calibration and to account for Galactic reddening (i.e., Figure 21), the GCM calibration was adjusted at the catalog level using SLR (Section 4.3). A reference stellar locus was empirically derived from the globally calibrated DES Y1A1 stellar objects in the region of the Y1A1 footprint with the smallest $E(B - V)$ value from Schlegel et al. (1998). Corrections were computed for the `WAVGCALIB_MAG_PSF` magnitudes described in Section 6. Our stellar selection was based on the weighted average of the `spread_model` quantity for the matched objects ($|\text{WAVG_SPREAD_MODEL_R}| < 0.003$). We selected co-add objects with $S/N > 10$ in i band and $S/N > 5$ in at least two other bands ($grzY$). We segmented the sky into equal-area pixels using the `HEALPix` scheme (Górski et al. 2005), starting with a relatively fine grid, $n_{\text{side}} = 512$ ($\sim 0.01 \text{ deg}^2$). If there were fewer than 200 stars in a pixel, then we appended neighboring pixels using the `get_all_neighbors` function from `healpy`, enlarging the pixel chunks until they contained at least 200 stars. To reduce computation time in high-density regions near the LMC, when there were more than 2000 stars per pixel we randomly down-sampled. Approximately 97% of the wide-area survey footprint was fit in chunks of 9 pixels containing a median of ~ 400 stars and yielding an effective resolution of $\sim 0.1 \text{ deg}^2$. We applied a modified version of the `BigMACS` SLR code (Kelly et al. 2014)⁸¹ to calibrate each star from the reference exposure with respect to the empirical stellar locus. The absolute calibration was set against the i -band

⁸¹ <https://code.google.com/p/big-macs-calibrate/>

magnitude derived from the GCM solution, which was dereddened using the SFD map with a reddening law of $A_I = 1.947 \times E(B - V)_{\text{SFD}}$. This extinction correction was derived following the prescription of Cardelli et al. (1989) with $R_V = 3.1$, but updated for the DES i -band throughput using optical–near-IR coefficients from O’Donnell (1994) assuming a source spectrum that is constant in spectral flux density per unit wavelength, f_λ (erg cm⁻² s⁻¹ Å⁻¹). The flat SED was chosen to represent the wide range of stellar SEDs from the Pickles ATLAS (Pickles 1998) and SEDs of galaxies over the range of redshifts probed by DES (Arnouts & Ilbert 2011).

Variations in the average metallicity of the stellar populations used for the SLR will introduce systematic shifts that are not due to photometric variation or Galactic reddening (e.g., High et al. 2009). For DES Y1A1, these shifts are largest for the g band, where they can have a 1%–2% effect on the calibration. A larger effect can be found in the vicinity of the LMC, which we avoid for extragalactic science. The effect of metallicity variations can be much worse at lower Galactic latitudes and in bluer filters (i.e., u band).

The final product was an SLR correction map at a resolution of $n_{\text{side}} = 512$ that we implemented with a bi-linear interpolation to obtain magnitude and flux corrections for the full Y1A1 GOLD catalog. The resulting SLR-adjusted magnitudes used in the Y1A1 GOLD catalog are thus already corrected for Galactic reddening.

A.4. Photometric Transformation Equations

We have derived transformation equations between various surveys and the DES system. We document these transformation equations here for reference.

We define a transformation from SDSS/UKIDSS to the DES system to place the nightly standard-star exposures on an initial DES AB photometric system (Section 4.1):

$$g_{\text{des}} = g_{\text{sdss}} - 0.104 \times (g - r)_{\text{sdss}} + 0.01, \quad (14)$$

$$r_{\text{des}} = r_{\text{sdss}} - 0.102 \times (g - r)_{\text{sdss}} + 0.02, \quad (15)$$

$$i_{\text{des}} = i_{\text{sdss}} - 0.256 \times (i - z)_{\text{sdss}} + 0.02, \quad (16)$$

$$z_{\text{des}} = z_{\text{sdss}} - 0.086 \times (i - z)_{\text{sdss}} + 0.01, \quad (17)$$

$$Y_{\text{des}} = Y_{\text{ukidss}} + 0.238 \times (z_{\text{sdss}} - Y_{\text{ukidss}}) + 0.634. \quad (18)$$

These transformation equations were derived in a hybrid manner: the color coefficients were determined by matching data from the DES SV data set with data from SDSS DR9 (or, in the case of the Y band, with a combination of UKIDSS DR6 Y band and SDSS DR9) and fitting the result. The zero-point for each relation was determined from synthetic AB photometry. We applied the DES, SDSS, and UKIDSS filter curves to the Pickles (1998) stellar library and measured the offset between the two synthetic magnitudes at zero color for each filter band. We note that the large zero-point offset for the Y -band transformation is due to the fact that the UKIDSS data are in the Vega magnitude system, while the Y_{des} is set to the AB magnitude system. These transformation equations are valid for stars with $(g - r)_{\text{sdss}} < 1.2$. For an individual object, the transformation from SDSS/UKIDSS to DES will depend on interstellar extinction. The DES footprint occupies a region of low extinction, and we estimate that the median correction due to reddening in the g band is 0.8 mmag (90% of Y1A1 GOLD has a g -band correction of < 2 mmag). Median

extinction corrections for the other bands are a factor of $\gtrsim 4$ lower than g band.

We validate the relative calibration accuracy of Y1A1 GOLD by comparing the calibrated magnitudes of stars in the Y1A1 GOLD catalog against those derived from a combination of APASS (Henden & Munari 2014) and 2MASS (Skrutskie et al. 2006). We selected stellar objects from the Y1A1 GOLD catalog using `MODEST_CLASS` (Section 8.1) and perform a 2'' match to the APASS and 2MASS catalogs. We then fit a set of transformation equations to map from g_{apass} , r_{apass} , and $J_{2\text{mass}}$ to a predicted magnitude in each of the DES filters:

$$g_{\text{des}} = g_{\text{apass}} - 0.0642 \times (g - r)_{\text{apass}} - 0.0239, \quad (19)$$

$$r_{\text{des}} = r_{\text{apass}} - 0.1264 \times (r - i)_{\text{apass}} - 0.0098, \quad (20)$$

$$i_{\text{des}} = r_{\text{apass}} - 0.4145 \times (r_{\text{apass}} - J_{2\text{mass}} - 0.81) - 0.391, \quad (21)$$

$$z_{\text{des}} = (J_{2\text{mass}} + 0.81) + 0.3866 \times (r_{\text{apass}} - J_{2\text{mass}} - 0.81) - 0.0414, \quad (22)$$

$$Y_{\text{des}} = (J_{2\text{mass}} + 0.81) + 0.2938 \times (r_{\text{apass}} - J_{2\text{mass}} - 0.81) - 0.0443. \quad (23)$$

Equations (21)–(23) are derived from a global fit of the Y1A1 GOLD data set and are valid for stars for which $r_{\text{apass}} - J_{2\text{mass}} < 1.81$. We find a cleaner and tighter relation using the hybrid APASS/2MASS ($r_{\text{apass}} - J_{2\text{mass}}$) color rather than a purely APASS ($r_{\text{apass}} - i_{\text{apass}}$) color for these transformation equations. These transformation equations explicitly remove any absolute calibration offset between the two data sets and can be used to test for spatial nonuniformity between the GCM calibration and these external catalogs (Figure 23). We note that the residual structure seen in the S82 region of Figure 23 does not appear in comparisons with SDSS DR10 or DES Y3, suggesting that this structure is a feature introduced by APASS.

To validate the completeness and contamination of the Y1A1 GOLD catalog, we perform a comparison with the CFHTLenS data in the W4 field. In this case, we are interested in the transformed magnitude of *all* objects, so we perform no stellar selection. We use matched objects to derive a set of transformation equations from the CFHTLenS g' , r' , i' , z' filters to the DES g , r , i , z system:

$$g_{\text{des}} = g_{\text{CFHT}} + 0.062(g_{\text{CFHT}} - r_{\text{CFHT}}) + 0.058, \quad (24)$$

$$r_{\text{des}} = r_{\text{CFHT}} - 0.078(g_{\text{CFHT}} - r_{\text{CFHT}}) + 0.021, \quad (25)$$

$$i_{\text{des}} = i_{\text{CFHT}} - 0.179(i_{\text{CFHT}} - z_{\text{CFHT}}) + 0.062, \quad (26)$$

$$z_{\text{des}} = z_{\text{CFHT}} - 0.139(i_{\text{CFHT}} - z_{\text{CFHT}}) + 0.053. \quad (27)$$

We find that these equations should be valid for objects with $g - r < 1.2$ and $i - z < 1.0$.

A.5. Calibration Validation

In this section we show ancillary plots of the performance and validation of the Y1A1 photometric calibration (Figures 22–25).

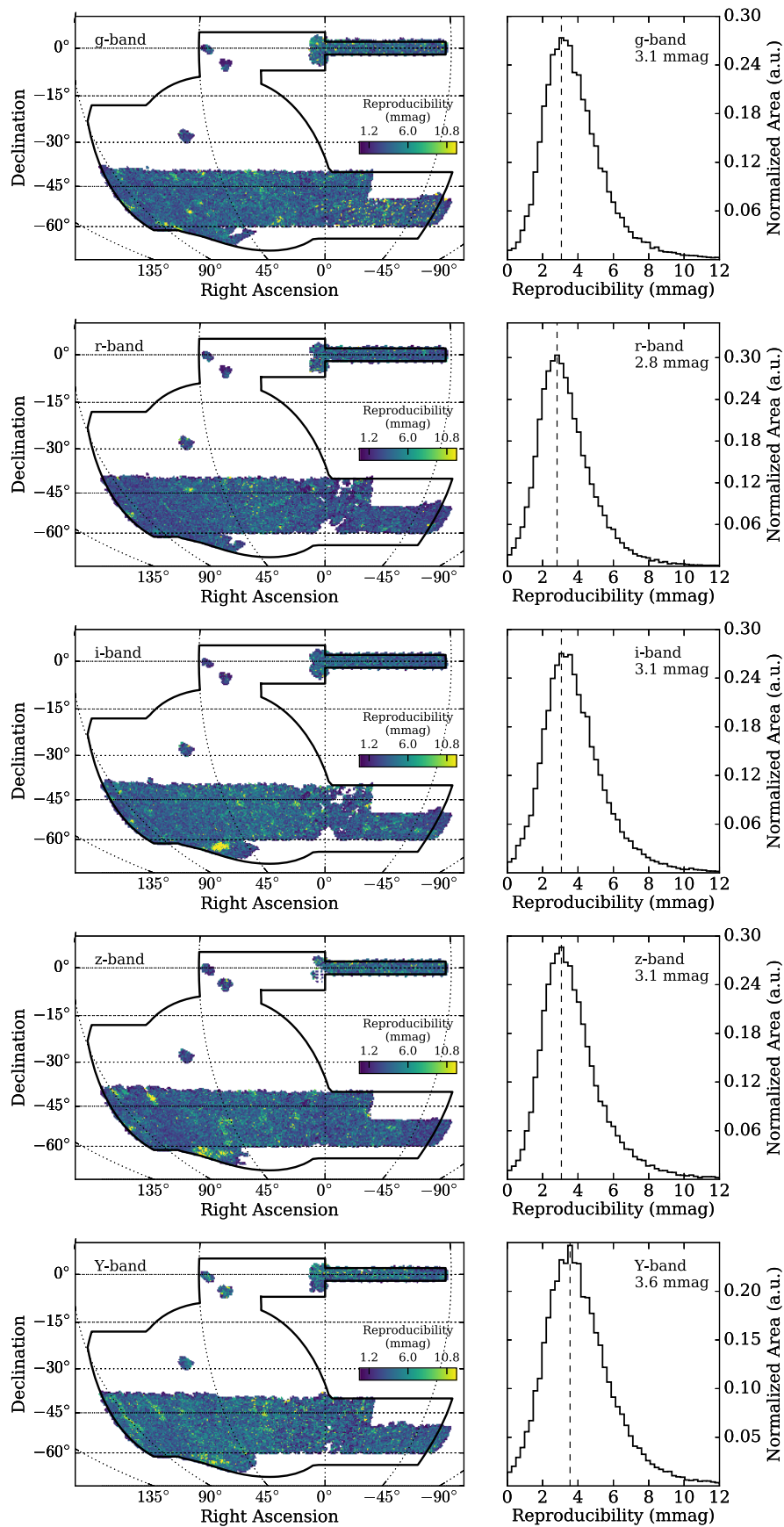


Figure 22. Internal rms errors in the photometric zero-point reproducibility per CCD for DES Y1A1. The zero-point rms is calculated by comparing the calibrated magnitudes of stars in overlapping CCDs. Note that these data include observations taken in both clear and cloudy conditions. Typical internal reproducibility errors are ~ 3 mmag ($\sim 0.3\%$). The color scale in the left panels represents the rms internal calibration uncertainty in mmag.

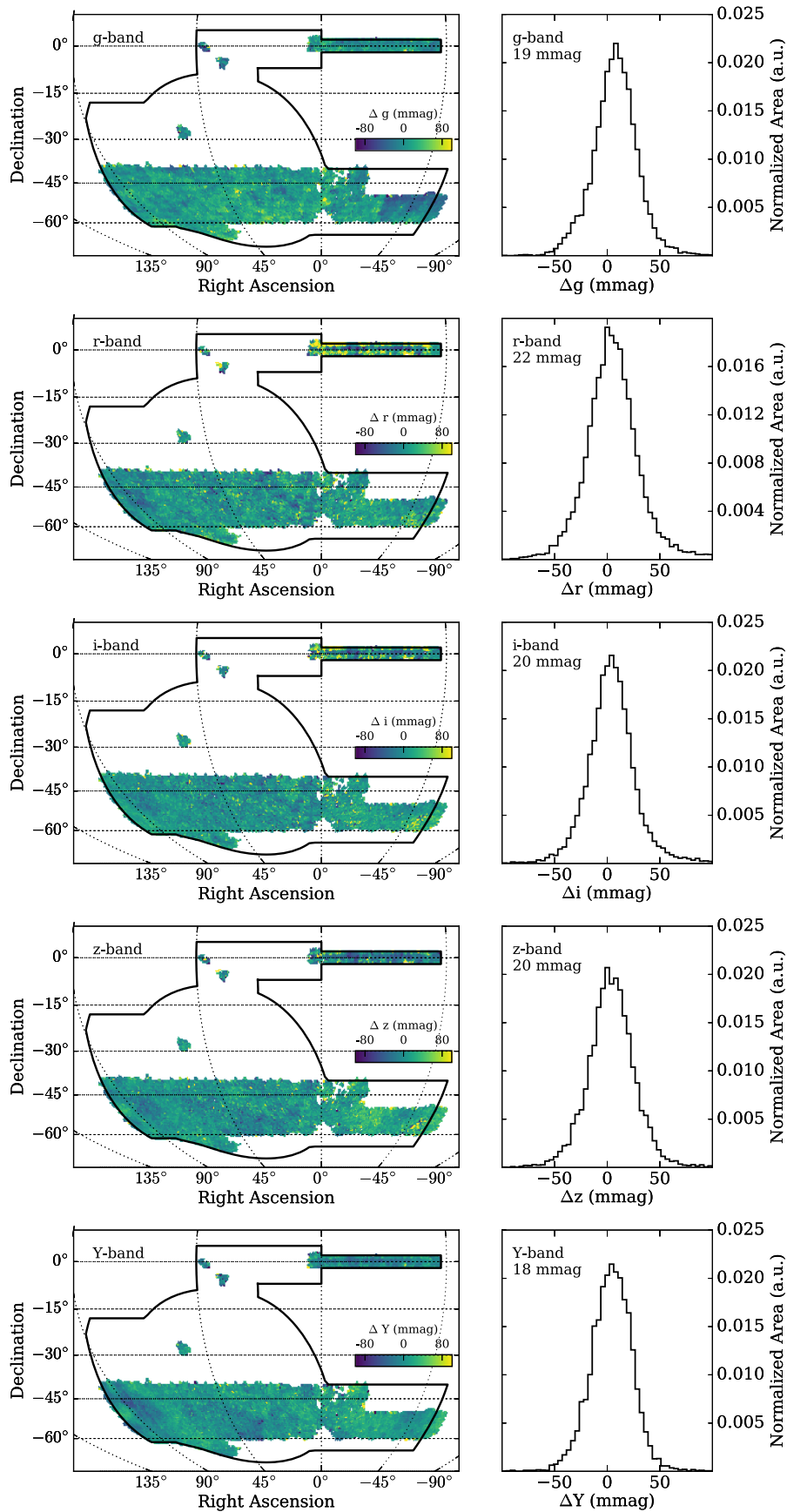


Figure 23. Comparison of stellar magnitudes from the DES Y1A1 GCM and those estimated from APASS/2MASS transformed into the DES filter system (Equations (19)–(23)). The sky plots (left) show the median magnitude offset for stars binned into $\sim 0.2 \text{ deg}^2$ HEALPIX pixels. The GCM calibrated magnitudes are consistent with the transformed values from APASS/2MASS with $\sigma_{68} \sim 20 \text{ mmag}$ (calculated between the 16th and 84th percentiles). Note that the GCM *g*-band calibration disagrees with APASS/2MASS by $\sim 4\%$ in the eastern portion of the SPT region ($\text{R.A.} < -20$), motivating the SLR adjustment described in Section 4.3.

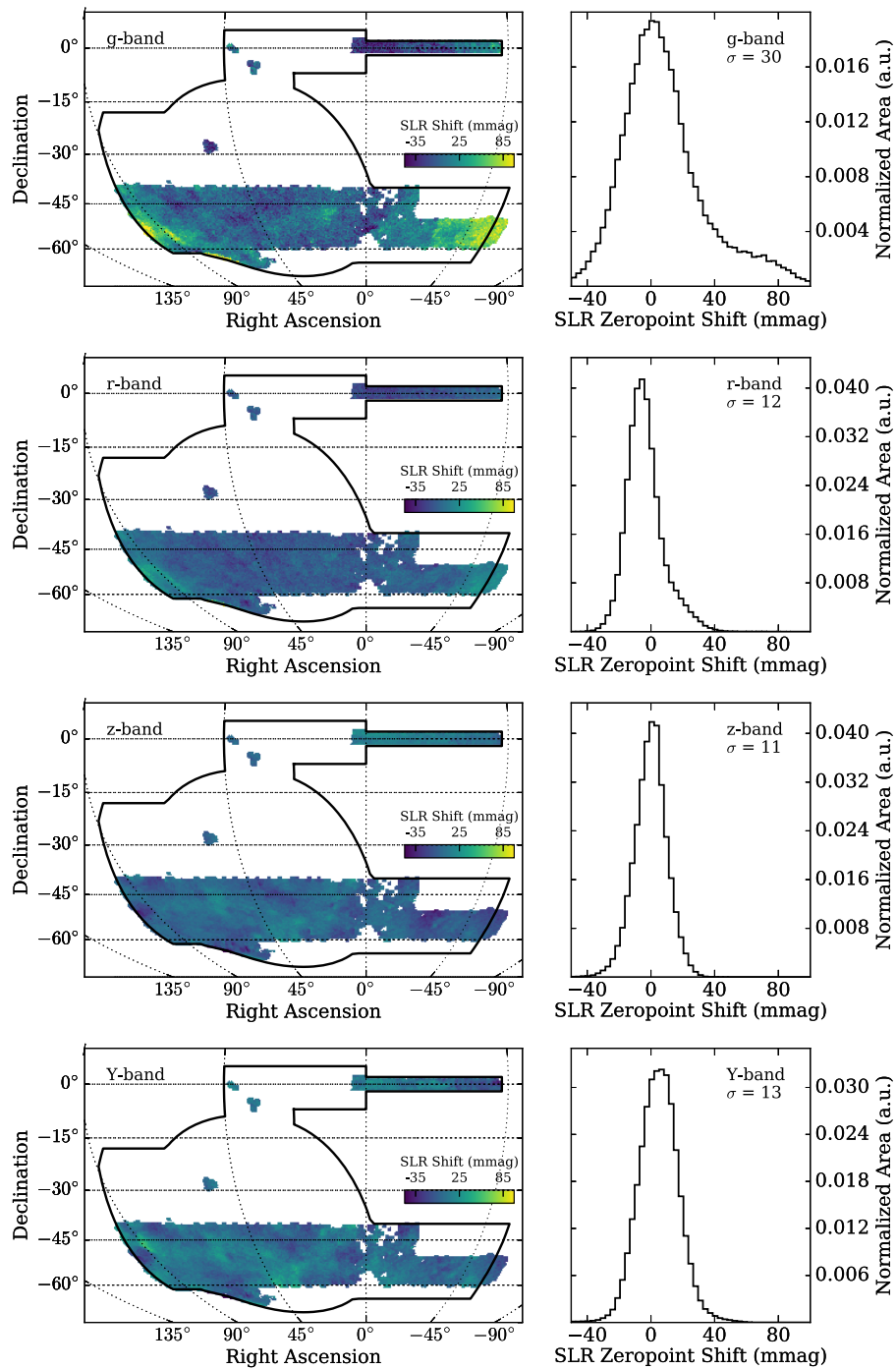


Figure 24. Adjustment to the GCM photometric zero-points derived from the SLR fit, after removing the contribution from interstellar extinction using the SFD maps and reddening from O’Donnell (1994). The width of these distributions represents adjustments to the calibration uniformity and differences between the interstellar extinction derived from the stellar locus and interstellar dust maps. The SLR adjustment is generally ~ 10 mmag (rms) over most of the area. A larger adjustment is made in the *g* band, which reflects the larger impact of reddening in the blue filters and a region of nonuniformity in the west of the footprint. There is no adjustment to the GCM *i* band because the SLR fit is tied to the dereddened magnitudes of stars in that band.

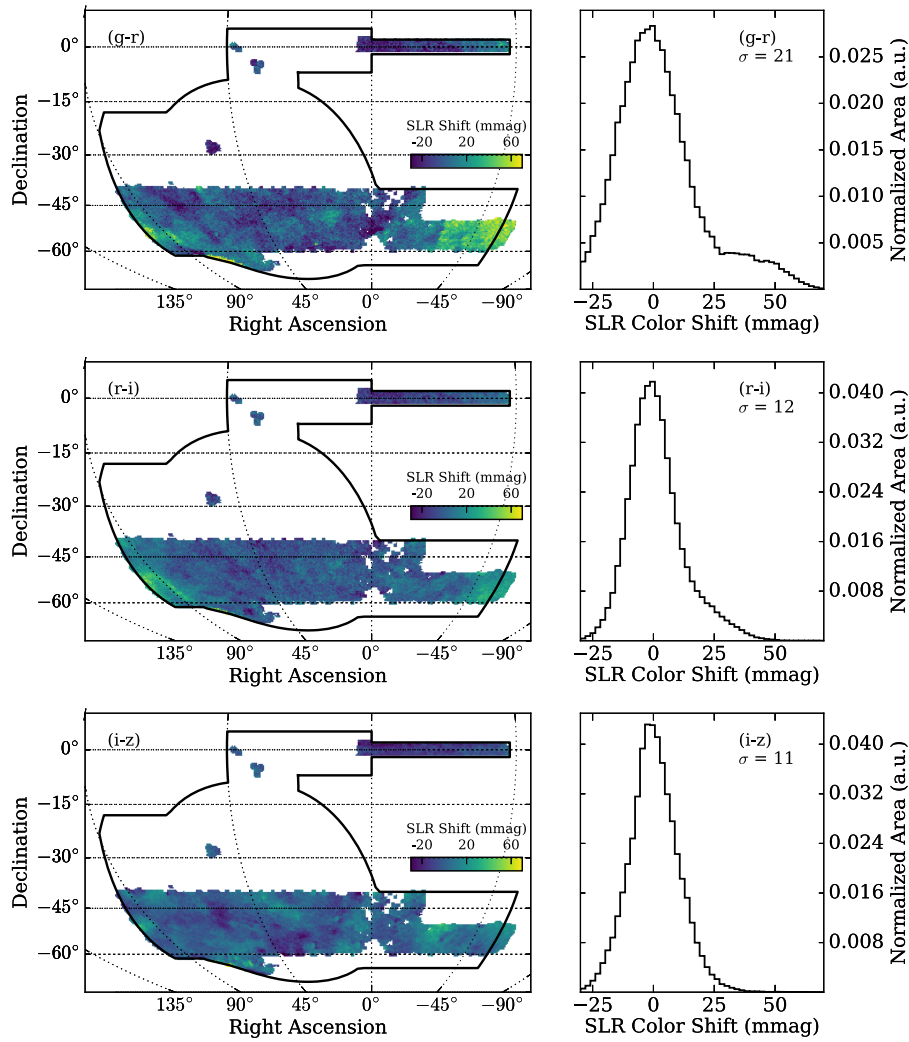


Figure 25. Color uniformity of the SLR adjustment applied to the GCM zero-points. The adjustment was largest for the $(g - r)$ color in the eastern portion of the SPT region. The color nonuniformity in this region was one of the specific motivations for the SLR calibration adjustment. After the SLR adjustment was applied, the color of stars was found to be uniform at the 1%–2% level across the footprint.

Appendix B COADD Source Detection

The Y1A1 COADD source detection was performed on a normalized “detection image” formed from a nonlinear combination of the r , i , and z co-added images. The original SWarp combination formula for computing the value of a pixel of the detection image is (Bertin 2010)

$$\chi = \sqrt{\frac{\sum_{c \leq n} w_c f_c^2}{n}}, \quad (28)$$

where f_c is the background-subtracted pixel value, w_c is the weight of the pixel in channel c , and n is the number of valid inputs. Compared to the standard χ^2 combination proposed by Szalay et al. (1999), χ leads to a less skewed noise distribution (if one assumes that input noise follows a Gaussian distribution), while maintaining identical detection capabilities. However, both estimators have a bias that depends on n , which leads to visible seams between regions with a different number of input images. This motivated the implementation of two new normalized image combination schemes in SWarp, with a variable offset applied to the original (still assuming that the

inputs are normally and independently distributed). CHI-MEAN is recentered on the mean (e.g., Evans et al. 2000):

$$\text{CHI-MEAN} = \frac{\sqrt{\sum_{c \leq n} w_c f_c^2} - \mu}{\sqrt{n - \mu^2}}, \quad (29)$$

with

$$\mu = \sqrt{2} \frac{\Gamma((n + 1)/2)}{\Gamma(n/2)}, \quad (30)$$

while CHI-MODE is recentered on the mode of the distribution:

$$\text{CHI-MODE} = \frac{\sqrt{\sum_{c \leq n} w_c f_c^2} - \sqrt{n - 1}}{\sqrt{n - \mu^2}}. \quad (31)$$

The left panel of Figure 26 shows a comparison of the distributions obtained from the original χ , CHI-MODE, and CHI-MEAN estimators for Gaussian input noise. The right panel of Figure 26 shows that the CHI-MEAN estimator generates the most seamless stacking results, and it was used to produce the Y1A1 COADD detection images.

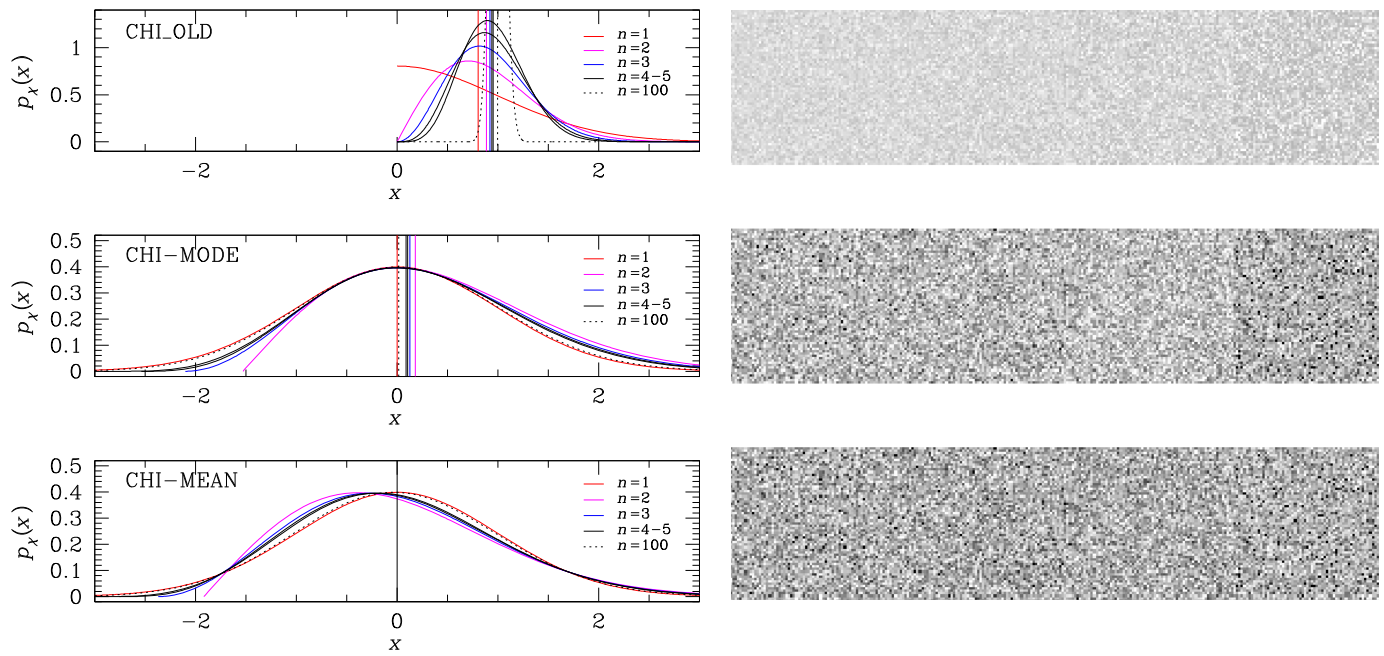


Figure 26. Left: normalized distribution of the value x of a detection image pixel for the original χ (OLD_CHI, top), CHI-MODE (middle), and CHI-MEAN (bottom) estimators when the inputs to the co-add are normally and independently distributed. n is the number of input images. The means of the distributions are shown as vertical lines. Right: gamma-corrected close-up of a χ (OLD_CHI, top), CHI-MODE (middle), and CHI-MEAN (bottom) detection image computed from a set of eight input images with zero-mean Gaussian white noise. The number of inputs decreases by steps of 64 pixels from left (eight inputs) to right (one input). CHI-MEAN detection images are virtually seamless, even at the transition between one and two input images.

Appendix C Catalog Depth Maps

In this appendix we collect a set of figures documenting the 10σ limiting magnitude of the Y1A1 GOLD catalog as

described in Section 7.1. We include depth maps both for the MAG_AUTO values derived from the co-added images (Figure 27) and for the CM_MAG values derived from multi-epoch, multiobject fitting (Figure 28).

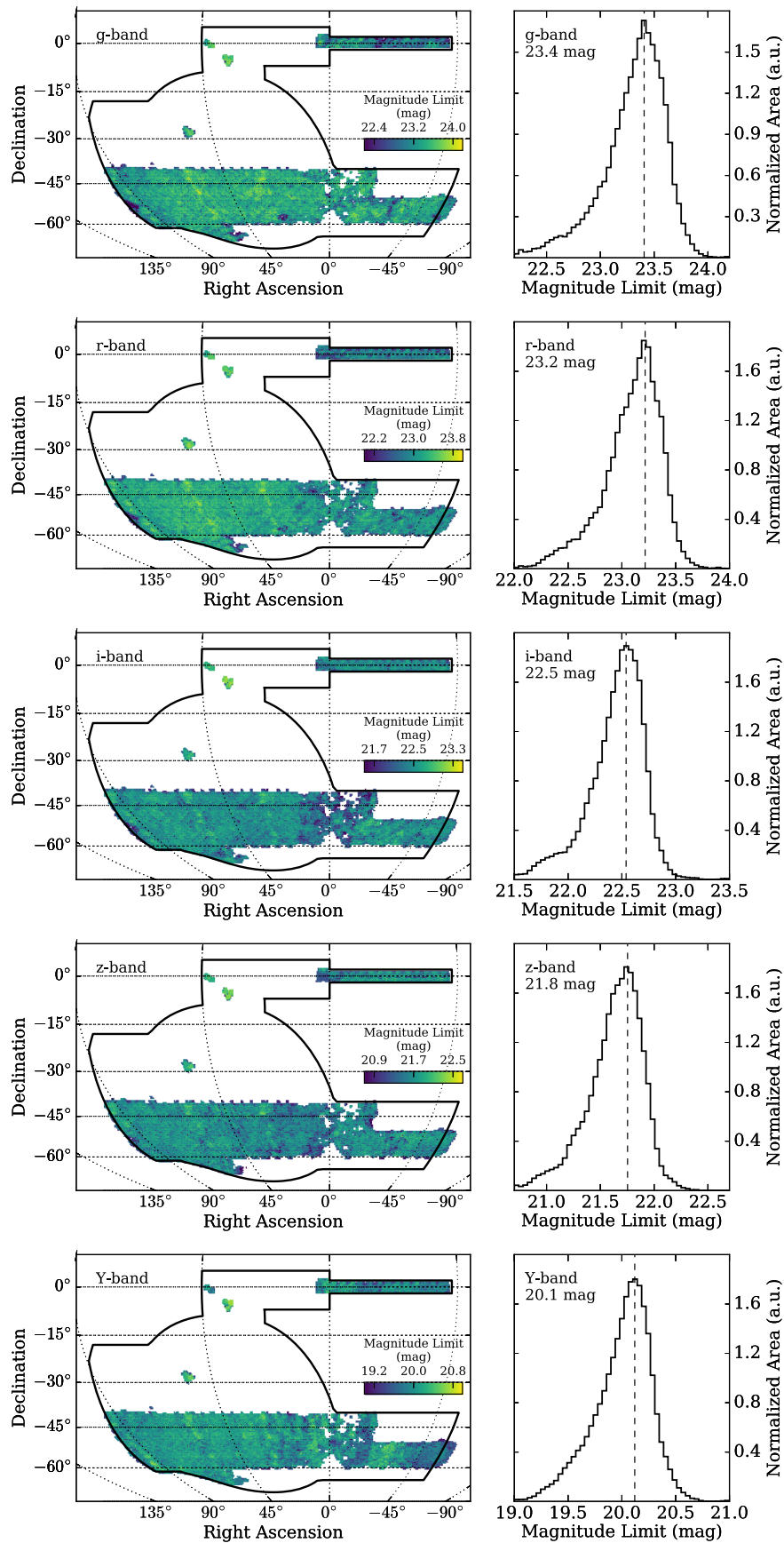


Figure 27. Sky maps and normalized histograms of the 10σ limiting magnitude for galaxies fit with `MAG_AUTO`. The mode of the limiting magnitude distribution is shown in the right panel of each row. The derivation of the limiting magnitude is described in more detail in Section 7.1.

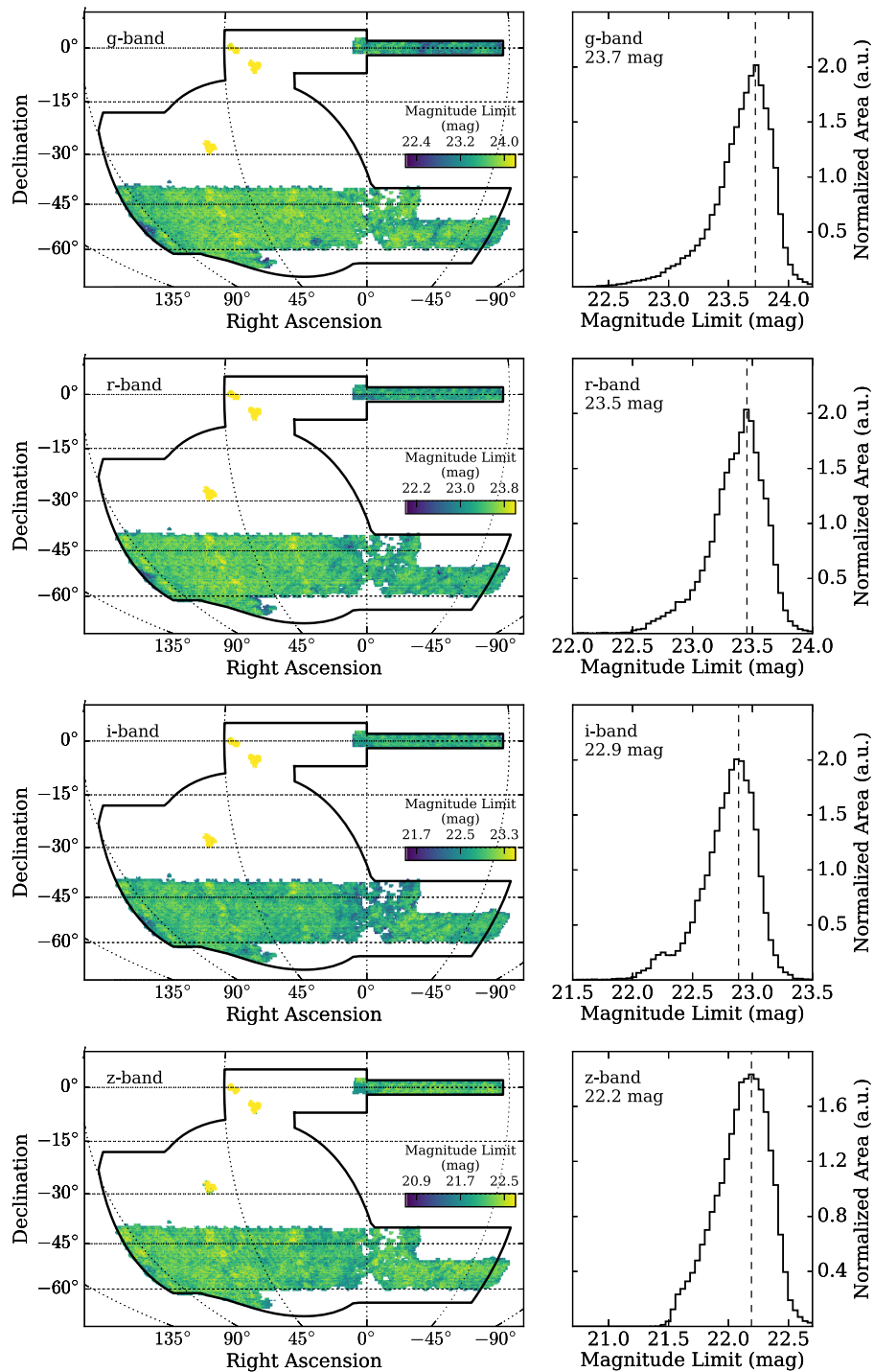


Figure 28. Sky maps and normalized histograms of the 10σ limiting magnitude for galaxies fit with the MOF CM_MAG . The magnitude range for these figures is the same as the MAG_AUTO magnitude limits shown in Figure 27. The mode of the limiting magnitude distribution is shown in the right panel of each row. Note that these magnitude limits include the deeper D10 co-adds of the SN fields. The derivation of the limiting magnitude is described in more detail in Section 7.1.

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References

- Ahn, C. P., Alexandroff, R., Allende Prieto, C., et al. 2012, *ApJS*, 203, 21
 Aihara, H., Armstrong, R., Bickerton, S., et al. 2017, arXiv:1702.08449
 Annis, J., Soares-Santos, M., Strauss, M. A., et al. 2014, *ApJ*, 794, 120
 Antilogus, P., Astier, P., Doherty, P., Guyonnet, A., & Regnault, N. 2014, *JInst*, 9, C03048
 Arnouts, S., & Ilbert, O. 2011, LePHARE: Photometric Analysis for Redshift Estimate, Astrophysics Source Code Library, ascl:1108.009
 Astropy Collaboration 2013, *A&A*, 558, A33
 Bechtol, K., Drlica-Wagner, A., Balbinot, E., et al. 2015, *ApJ*, 807, 50
 Benéz, N. 2000, *ApJ*, 536, 571
 Berk, A., Anderson, G. P., Bernstein, L. S., et al. 1999, *Proc. SPIE*, 3756, 348
 Berlind, A. A., Frieman, J., Weinberg, D. H., et al. 2006, *ApJS*, 167, 1
 Bernstein, G., Abbott, T., Desai, S., et al. 2017a, arXiv:1706.09928
 Bernstein, G. M., Armstrong, R., Plazas, A. A., et al. 2017b, *PASP*, 129, 074503
 Bertin, E. 2006, in ASP Conf. Ser. 351, Astronomical Data Analysis Software and Systems XV, ed. C. Gabriel et al. (San Francisco, CA: ASP), 112
 Bertin, E. 2010, SWarp: Resampling and Co-adding FITS Images Together, Astrophysics Source Code Library, ascl:1010.068
 Bertin, E. 2011, in ASP Conf. Ser. 442, Astronomical Data Analysis Software and Systems XX, ed. I. N. Evans et al. (San Francisco, CA: ASP), 435
 Bertin, E., & Arnouts, S. 1996, *A&AS*, 117, 393
 Bertin, E., Mellier, Y., Radovich, M., et al. 2002, in ASP Conf. Ser. 281, Astronomical Data Analysis Software and Systems XI, ed. D. A. Bohlender, D. Durand, & T. H. Handley (San Francisco, CA: ASP), 228
 Bohlin, R. C., Gordon, K. D., & Tremblay, P.-E. 2014, *PASP*, 126, 711
 Bonnett, C., Troxel, M. A., Hartley, W., et al. 2016, *PhRvD*, 94, 042005
 Bouy, H., Bertin, E., Moraux, E., et al. 2013, *A&A*, 554, A101
 Burke, D., Rykoff, E. S., Allam, S., et al. 2017, arXiv:1706.01542
 Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245
 Carlstrom, J. E., Ade, P. A. R., Aird, K. A., et al. 2011, *PASP*, 123, 568
 Casali, M., Adamson, A., Alves de Oliveira, C., et al. 2007, *A&A*, 467, 777
 Cawthon, R., Davis, C., Gatti, M., et al. 2017, arXiv:1712.07298
 Chang, C., Busha, M. T., Wechsler, R. H., et al. 2015, *ApJ*, 801, 73
 Corwin, H. G., Jr., Buta, R. J., & de Vaucouleurs, G. 1994, *AJ*, 108, 2128
 Coupon, J., Kilbinger, M., McCracken, H. J., et al. 2012, *A&A*, 542, A5
 Davis, C., Gatti, M., Vielzeuf, P., et al. 2017, PRD, submitted
 de Vaucouleurs, G. 1948, *AnAp*, 11, 247
 De Vicente, J., Sánchez, E., & Sevilla-Noarbe, I. 2016, *MNRAS*, 459, 3078
 DES Collaboration 2005, arXiv:astro-ph/0510346
 DES Collaboration 2016, *MNRAS*, 460, 1270
 DES Collaboration 2018, arXiv:1801.03181
 Desai, S., Armstrong, R., Mohr, J. J., et al. 2012, *ApJ*, 757, 83
 Diehl, H. T., Abbott, T. M. C., Annis, J., et al. 2014, *Proc. SPIE*, 9149, 91490V
 Drlica-Wagner, A., Bechtol, K., Rykoff, E. S., et al. 2015, *ApJ*, 813, 109
 Erben, T., Hildebrandt, H., Miller, L., et al. 2013, *MNRAS*, 433, 2545
 Evans, M., Hastings, N., & Peacock, B. 2000, *Statistical Distributions*, Wiley Series in Probability and Statistics (New York: Wiley)
 Fadely, R., Hogg, D. W., & Willman, B. 2012, *ApJ*, 760, 15
 Flaugher, B., Diehl, H. T., Honscheid, K., et al. 2015, *AJ*, 150, 150
 Gaia Collaboration 2016, *A&A*, 595, A2
 Gatti, M., Vielzeuf, P., Davis, C., et al. 2018, *MNRAS*, in press (arXiv:1709.00992)
 Giavalisco, M., Ferguson, H. C., Koekemoer, A. M., et al. 2004, *ApJL*, 600, L93
 Gilbank, D. G., Gladders, M. D., Yee, H. K. C., & Hsieh, B. C. 2011, *AJ*, 141, 94
 Glazebrook, K., Peacock, J. A., Collins, C. A., & Miller, L. 1994, *MNRAS*, 266, 65
 Górski, K. M., Hivon, E., Banday, A. J., et al. 2005, *ApJ*, 622, 759
 Gruen, D., Bernstein, G. M., Jarvis, M., et al. 2015, *JInst*, 10, C05032
 Gschwend, J., Rosell, A. C., Ogando, R., et al. 2017, *A&C*, submitted
 Hambly, N. C., Collins, R. S., Cross, N. J. G., et al. 2008, *MNRAS*, 384, 637
 Hamilton, A. J. S., & Tegmark, M. 2004, *MNRAS*, 349, 115
 Henden, A., & Munari, U. 2014, *CoSka*, 43, 518
 Hennig, C., Mohr, J. J., Zenteno, A., et al. 2017, *MNRAS*, 467, 4015
 Hewett, P. C., Warren, S. J., Leggett, S. K., et al. 2006, *MNRAS*, 367, 454
 Heymans, C., Van Waerbeke, L., Miller, L., et al. 2012, *MNRAS*, 427, 146
 High, F. W., Stubbs, C. W., Rest, A., Stalder, B., & Challis, P. 2009, *AJ*, 138, 110
 Hildebrandt, H., Erben, T., Kuijken, K., et al. 2012, *MNRAS*, 421, 2355
 Hodgkin, S. T., Irwin, M. J., Hewett, P. C., et al. 2009, *MNRAS*, 394, 675
 Hoffleit, D., & Jaschek, C. (ed.) 1991, *The Bright Star Catalogue* (New Haven, CT: Yale University Observatory)
 Hogg, D. W., & Lang, D. 2013, *PASP*, 125, 719
 Hoyle, B., Gruen, D., Bernstein, G. M., et al. 2017, *MNRAS*, submitted (arXiv:1708.01532)
 Huchra, J. P., & Geller, M. J. 1982, *ApJ*, 257, 423
 Hunter, J. D. 2007, *CSE*, 9, 90
 Ivezić, Ž., Lupton, R. H., Schlegel, D., et al. 2004, *AN*, 325, 583
 Jarvis, M., Sheldon, E., Zuntz, J., et al. 2016, *MNRAS*, 460, 2245
 Jones, E., Oliphant, T., Peterson, P., et al. 2001, *SciPy: Open Source Scientific Tools for Python*
 Kelly, P. L., von der Linden, A., Applegate, D. E., et al. 2014, *MNRAS*, 439, 28
 Kessler, R., Marriner, J., Childress, M., et al. 2015, *AJ*, 150, 172
 Klein, M., Mohr, J. J., Desai, S., et al. 2017, arXiv:1706.06577
 Kron, R. G. 1980, *ApJS*, 43, 305
 Laigle, C., McCracken, H. J., Ilbert, O., et al. 2016, *ApJS*, 224, 24
 Lawrence, A., Warren, S. J., Almaini, O., et al. 2007, *MNRAS*, 379, 1599
 Le Fèvre, O., Vettolani, G., Garilli, B., et al. 2005, *A&A*, 439, 845
 Leauthaud, A., Massey, R., Kneib, J.-P., et al. 2007, *ApJS*, 172, 219
 Leistedt, B., Peiris, H. V., Elsner, F., et al. 2016, *ApJS*, 226, 24
 Lewis, P. M., Rogers, H., & Schindler, R. H. 2010, *Proc. SPIE*, 7735, 77353C
 Li, T. S., DePoy, D. L., Marshall, J. L., et al. 2016, *AJ*, 151, 157
 MacDonald, E. C., Allen, P., Dalton, G., et al. 2004, *MNRAS*, 352, 1255
 Melchior, P., Sheldon, E., Drlica-Wagner, A., et al. 2016, *A&C*, 16, 99
 Ménard, B., Scranton, R., Schmidt, S., et al. 2013, arXiv:1303.4722
 Mohr, J. J., Adams, D., Barkhouse, W., et al. 2008, *SPIE*, 7016, 0
 Mohr, J. J., Armstrong, R., Bertin, E., et al. 2012, *Proc. SPIE*, 8451, 84510D
 Morganson, E., Gruendl, R., & Kent, S. 2018, *PASP*, in press (arXiv:1801.03177)
 Neilsen, E., Bernstein, G., Gruendl, R., & Kent, S. 2015, "Limiting Magnitude, τ , Teff, and Image Quality in DES Year 1," Tech. Rep. FERMILAB-TM-2610-AE-CD
 Newman, J. A. 2008, *ApJ*, 684, 88
 Ngeow, C., Mohr, J. J., Alam, T., et al. 2006, *Proc. SPIE*, 6270, 627023
 O'Donnell, J. E. 1994, *ApJ*, 422, 158
 Padmanabhan, N., Schlegel, D. J., Finkbeiner, D. P., et al. 2008, *ApJ*, 674, 1217
 Pickles, A. J. 1998, *PASP*, 110, 863
 Plazas, A. A., Bernstein, G. M., & Sheldon, E. S. 2014, *PASP*, 126, 750
 Reil, K., Lewis, P., Schindler, R., & Zhang, Z. 2014, *Proc. SPIE*, 9149, 91490U
 Rozo, E., Rykoff, E. S., Abate, A., et al. 2016, *MNRAS*, 461, 1431
 Rykoff, E. S., Rozo, E., & Keisler, R. 2015, arXiv:1509.00870
 Sánchez, C., Clampitt, J., Kovacs, A., et al. 2017, *MNRAS*, 465, 746
 Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, *ApJ*, 500, 525
 Schmidt, S. J., Ménard, B., Scranton, R., et al. 2013, *MNRAS*, 431, 3307
 Scoville, N., Aussel, H., Brusa, M., et al. 2007, *ApJS*, 172, 1
 Sevilla, I., Armstrong, R., Bertin, E., et al. 2011, arXiv:1109.6741
 Sheldon, E. S. 2014, *MNRAS*, 444, L25
 Sheldon, E. S., & Huff, E. M. 2017, *ApJ*, 841, 24
 Sinnott, R. W. 1988, in *NGC 2000.0: The Complete New General Catalogue and Index Catalogues of Nebulae and Star Clusters*, ed. J. L. E. Dreyer (Cambridge: Cambridge Univ. Press)
 Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, 131, 1163
 Smith, J. A., Allam, S. S., Tucker, D. L., et al. 2017, *AJ*, submitted
 Smith, J. A., Tucker, D. L., Kent, S., et al. 2002, *AJ*, 123, 2121
 Soumagnac, M. T., Abdalla, F. B., Lahav, O., et al. 2015, *MNRAS*, 450, 666
 Suchyta, E., Huff, E. M., Aleksić, J., et al. 2016, *MNRAS*, 457, 786
 Swanson, M. E. C., Tegmark, M., Hamilton, A. J. S., & Hill, J. C. 2008, *MNRAS*, 387, 1391
 Szalay, A. S., Connolly, A. J., & Szokoly, G. P. 1999, *AJ*, 117, 68
 Tucker, D. L., Annis, J. T., Lin, H., et al. 2007, in ASP Conf. Ser. 364, *The Future of Photometric, Spectrophotometric and Polarimetric Standardization*, ed. C. Sterken (San Francisco, CA: ASP), 187
 Tucker, D. L., Kent, S., Richmond, M. W., et al. 2006, *AN*, 327, 821

Van Der Walt, S., Colbert, S. C., & Varoquaux, G. 2011, [CSE](#), 13, 22
Wittman, D. 2009, [ApJL](#), 700, L174
Wyatt, S., Tucker, D. L., & Smith, A. 2014, American Astronomical Society Meeting Abstracts, 223, [254.13](#)

Zacharias, N., Finch, C. T., Girard, T. M., et al. 2013, [AJ](#), 145, 44
Zhang, Y., McKay, T. A., Bertin, E., et al. 2015, [PASP](#), 127, 1183
Zuntz, J., Sheldon, E. S., Samuroff, S., et al. 2017, arXiv:[1708.01533](#)