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Jason T. Hinkle ¹★, T. W.-S. Holoién ²†, K. Auchettl ^{3,4,5,6}, B. J. Shappee ¹,
 J. M. M. Neustadt ⁷, A. V. Payne ¹‡, J. S. Brown,⁴ C. S. Kochanek ^{7,8}, K. Z. Stanek,^{7,8}
 M. J. Graham ⁹, M. A. Tucker ¹§, A. Do,¹ J. P. Anderson ¹⁰, S. Bose ^{7,8}, P. Chen,¹¹ D. A. Coulter,⁴
 G. Dimitriadis,⁴ Subo Dong ¹¹, R. J. Foley,⁴ M. E. Huber ¹, T. Hung,⁴ C. D. Kilpatrick,⁴ G. Pignata,^{12,13}
 A. L. Piro,² C. Rojas-Bravo ⁴, M. R. Siebert,⁴ B. Stalder,¹⁴ Todd A. Thompson ^{7,8}, J. L. Tonry,¹
 P. J. Vallely⁷ and J. P. Wisniewski ¹⁵

Affiliations are listed at the end of the paper

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ABSTRACT

We present observations of ASASSN-19dj, a nearby tidal disruption event (TDE) discovered in the post-starburst galaxy KUG 0810+227 by the All-Sky Automated Survey for Supernovae (ASAS-SN) at a distance of $d \simeq 98$ Mpc. We observed ASASSN-19dj from -21 to 392 d relative to peak ultraviolet (UV)/optical emission using high-cadence, multiwavelength spectroscopy and photometry. From the ASAS-SN g -band data, we determine that the TDE began to brighten on 2019 February 6.8 and for the first 16 d the rise was consistent with a flux $\propto t^2$ power law. ASASSN-19dj peaked in the UV/optical on 2019 March 6.5 (MJD = 58548.5) at a bolometric luminosity of $L = (6.2 \pm 0.2) \times 10^{44}$ erg s⁻¹. Initially remaining roughly constant in X-rays and slowly fading in the UV/optical, the X-ray flux increased by over an order of magnitude ~ 225 d after peak, resulting from the expansion of the X-ray emitting region. The late-time X-ray emission is well fitted by a blackbody with an effective radius of $\sim 1 \times 10^{12}$ cm and a temperature of $\sim 6 \times 10^5$ K. The X-ray hardness ratio becomes softer after brightening and then returns to a harder state as the X-rays fade. Analysis of Catalina Real-Time Transient Survey images reveals a nuclear outburst roughly 14.5 yr earlier with a smooth decline and a luminosity of $L_V \geq 1.4 \times 10^{43}$ erg s⁻¹, although the nature of the flare is unknown. ASASSN-19dj occurred in the most extreme post-starburst galaxy yet to host a TDE, with Lick $H\delta_A = 7.67 \pm 0.17$ Å.

Key words: accretion, accretion discs – black hole physics – galaxies: nuclei.

1 INTRODUCTION

Supermassive black holes (SMBHs) are known to reside in the centres of most massive galaxies (e.g. Rees 1988; Kormendy & Richstone 1995; Magorrian et al. 1998; Ho 2008; Gültekin et al. 2009; Kormendy & Ho 2013). If mass is actively accreting on to these SMBHs, they can be detected as active galactic nuclei (AGNs). Conversely, direct detections of inactive SMBHs are difficult, mainly limited to our own black hole (Sgr A*); Ghez et al. 2005), or massive ($\gtrsim 10^6 M_\odot$) SMBHs in nearby ($\lesssim 50$ Mpc) galaxies, where stars (e.g. Kormendy et al. 1996; Gebhardt et al. 2011) and/or gas (e.g. Ford et al. 1994; Atkinson et al. 2005) within the SMBH's sphere of influence can be resolved. Only one SMBH, Pōwehi in M87, has been directly observed, by the Event Horizon Telescope (Event Horizon Telescope Collaboration et al. 2019). Tidal disruption events (TDEs) provide an opportunity to study otherwise inactive SMBHs at greater distances.

A TDE occurs when a star passes within the tidal radius of an SMBH and is torn apart, resulting in a luminous accretion flare (Rees 1988; Evans & Kochanek 1989; Phinney 1989; Ulmer 1999; Komossa 2015; Stone et al. 2019). Early theoretical work predicted that the blackbody temperatures of TDEs should be of the order of 10^5 K, consistent with a peak in the soft X-ray band (e.g. Lacy, Townes & Hollenbach 1982; Rees 1988; Evans & Kochanek 1989; Phinney 1989), but observational studies have discovered a breadth of TDE phenomenology. For example, TDE candidates have been detected in the hard X-ray (e.g. Bloom et al. 2011; Burrows et al. 2011; Cenko et al. 2012b; Pasham et al. 2015), soft X-ray (e.g. Bade, Komossa & Dahlem 1996; Grupe, Thomas & Leighly 1999; Komossa & Bade 1999; Komossa & Greiner 1999; Auchettl, Guillochon & Ramirez-Ruiz 2017), ultraviolet (UV) (e.g. Stern et al. 2004; Gezari et al. 2006, 2008, 2009), optical (e.g. van Velzen et al. 2011; Cenko et al. 2012a; Gezari et al. 2012; Arcavi et al. 2014; Chornock et al. 2014; Holoién et al. 2014b; Vinkó et al. 2015; Holoién et al. 2016b, c; Brown et al. 2018; Holoién et al. 2019a, b), and radio (e.g. Zauderer et al. 2011; Cenko et al. 2012b; Alexander et al. 2016, 2017; van Velzen et al. 2016; Brown et al. 2017), with many showing emission in multiple energy bands.

The diversity seen in these events fuelled a broad range of theoretical investigations (e.g. Lodato et al. 2015; Krolik et al.

* E-mail: jhinkle6@hawaii.edu

† Carnegie Fellow.

‡ NASA Fellowship Activity Fellow.

§ DOE CSGF Fellow.

2016; Svirski, Piran & Krolik 2017; Krolik, Piran & Ryu 2020; Ryu et al. 2020a). The unifying model of Dai et al. (2018) may provide an explanation of the diversity, positing that many of the observed multiwavelength photometric and spectroscopic properties of TDEs are a result of the viewing angle. Despite this, the origin of the UV/optical emission is still debated, with reprocessed emission from an accretion disc (e.g. Dai et al. 2018; Mockler, Guillochon & Ramirez-Ruiz 2019) and shocks from stream-stream collisions (e.g. Jiang, Guillochon & Loeb 2016; Bonnerot, Rossi & Lodato 2017; Lu & Bonnerot 2020; Ryu et al. 2020a) as the most commonly proposed emission mechanisms. Due to similar energy budgets and the general lack of observational features to neatly disambiguate models for a particular event, the nature of the UV/optical emission remains an open question.

Observations of TDEs may provide information on the physics of accretion (e.g. Lodato & Rossi 2011; Guillochon & Ramirez-Ruiz 2015; Shiokawa et al. 2015; Metzger & Stone 2016), shock physics (e.g. Lodato, King & Pringle 2009), jet formation (e.g. Farrar & Piran 2014; Wang & Liu 2016; Biehl et al. 2018), and the environment and growth of SMBHs (e.g. Auchettl, Ramirez-Ruiz & Guillochon 2018). However, the characteristics of the observed emission from TDEs, such as their light curves, spectroscopic evolution (both optical and X-ray), blackbody properties, etc., are a function of many physical parameters. Such properties include the star's impact parameter (e.g. Guillochon & Ramirez-Ruiz 2013, 2015; Gafton & Rosswog 2019), mass (e.g. Gallegos-Garcia, Law-Smith & Ramirez-Ruiz 2018; Law-Smith, Guillochon & Ramirez-Ruiz 2019; Mockler et al. 2019), composition (e.g. Kochanek 2016a), evolutionary stage (e.g. MacLeod et al. 2012), age (e.g. Gallegos-Garcia et al. 2018), and spin (e.g. Golightly, Coughlin & Nixon 2019). Additionally, stellar demographics (e.g. Kochanek 2016b), the fraction of accreted stellar material (e.g. Metzger & Stone 2016; Coughlin & Nixon 2019), and the geometry of accretion (e.g. Kochanek 1994; Lodato & Rossi 2011; Dai, McKinney & Miller 2015; Guillochon & Ramirez-Ruiz 2015; Shiokawa et al. 2015; Metzger & Stone 2016; Dai et al. 2018) may affect the observed emission.

It has also been shown that TDE emission may be sensitive to black hole spin and mass (e.g. Ulmer 1999; Graham et al. 2001; Gafton & Rosswog 2019; Mockler et al. 2019), making TDEs useful probes of otherwise quiescent SMBHs. As such, TDE light curves can be used to constrain the masses of SMBHs, which are consistent with those derived from other methods (Mockler et al. 2019). While there are a large number of potentially relevant physical parameters, the observed UV/optical emission is relatively well fitted by a blackbody (e.g. Gezari et al. 2012; Holoien et al. 2014a, 2016b, c; Brown et al. 2016; Hung et al. 2017; Holoien et al. 2018, 2019b; Leloudas et al. 2019; Holoien et al. 2020; van Velzen et al. 2020). It has also been shown that the peak UV/optical luminosities of TDEs are related to their decline rates (Mockler et al. 2019; Hinkle et al. 2020), with more luminous TDEs declining more slowly after peak. As the number of TDEs increases, they will provide a more complete picture of SMBH growth and evolution via accretion and the central environments of galaxies.

The spectroscopic properties of optical TDEs are varied (e.g. Arcavi et al. 2014; Hung et al. 2017; Leloudas et al. 2019; Wevers et al. 2019b; Holoien et al. 2020; van Velzen et al. 2020), with differences in observed species, line strengths/widths, and line ratios. Emission lines from hydrogen, helium, and more exotic features due to Bowen fluorescence have been observed (e.g. Leloudas et al. 2019; van Velzen et al. 2020). Most TDEs have simple broad-line profiles, but others have double-peaked disc-like line profiles (e.g. Holoien et al. 2019a; Hung et al. 2020) or strong narrow lines (e.g.

Holoien et al. 2020; van Velzen et al. 2020). Possible explanations for this variety are details in the physics of photoionization (e.g. Gaskell & Rojas Lobos 2014; Guillochon, Manukian & Ramirez-Ruiz 2014; Roth et al. 2016; Kara et al. 2018; Leloudas et al. 2019), differences in the composition of stars due to evolution (Kochanek 2016a), the viewing geometry with respect to an accretion disc (e.g. Holoien et al. 2019a; Hung et al. 2020; Short et al. 2020), or even the disruption of helium stars (Gezari et al. 2012; Strubbe & Murray 2015). Additionally, there have been variations in the times at which strong emission lines appear (e.g. Holoien et al. 2019a, b, 2020). An even larger sample of optically bright TDEs will better constrain the mechanisms that influence the observed emission from such events.

TDEs are rare, with an expected frequency between 10^{-4} and 10^{-5} yr $^{-1}$ per galaxy (e.g. van Velzen & Farrar 2014; Holoien et al. 2016b; Auchettl et al. 2018; van Velzen 2018). Interestingly though, TDEs seem to prefer post-starburst host galaxies. In such galaxies, the TDE rates can be enhanced by up to 200 times as compared to the average rates (e.g. Arcavi et al. 2014; French, Arcavi & Zabludoff 2016; Law-Smith et al. 2017; Graur et al. 2018). Combining these suggests that in the most extreme post-starbursts, TDEs can occur at roughly the same rate as other bright transients like supernovae.

There are few observations of the early-time evolution of TDEs, a time period that may be important to understanding how the disrupted stellar material settles into an accretion flow. With the advent of transient surveys like the All-Sky Automated Survey for Supernovae (ASAS-SN; Shappee et al. 2014; Kochanek et al. 2017), the Asteroid Terrestrial Impact Last Alert System (ATLAS; Tonry et al. 2018), the Zwicky Transient Facility (ZTF; Bellm et al. 2019), the Panoramic Survey Telescope and Rapid Response System (Pan-STARRS; Chambers et al. 2016), and the Young Supernova Experiment (YSE; Jones et al. 2019), many more TDEs are being discovered. This includes an increasing number of TDEs discovered before their peak brightness (e.g. Holoien et al. 2019a, b; Leloudas et al. 2019; van Velzen et al. 2019; Wevers et al. 2019b; Holoien et al. 2020; van Velzen et al. 2020).

While these fast-cadence, wide-field optical surveys are ideal for discovering TDEs, a significant fraction of emission from some events is in the soft X-ray band (e.g. Ulmer 1999; Auchettl et al. 2017). Recently, an increasing number of TDE candidates discovered in the optical have exhibited strong X-ray emission. Examples include ASASSN-14li (e.g. Miller et al. 2015; Holoien et al. 2016b; Brown et al. 2017), ASASSN-15oi (e.g. Holoien et al. 2016a; Gezari, Cenko & Arcavi 2017; Holoien et al. 2018), ASASSN-18ul (Wevers et al. 2019b; Payne et al., in preparation), Gaia19bpt (van Velzen et al. 2020), ZTF19aapreis (van Velzen et al. 2020), and the TDE studied in this work, ASASSN-19dj. The combination of UV, optical, and X-ray has given greater insight on the formation of an accretion disc, reprocessing, and the differences between thermal (non-jetted) and non-thermal (jetted) TDEs (e.g. Auchettl et al. 2017). In general, long-term X-ray light curves of TDE candidates are required to distinguish them from AGNs and to study detailed accretion physics (e.g. Auchettl et al. 2018).

In this paper, we present the discovery and observations of ASASSN-19dj. Smaller data sets on ASASSN-19dj have been analysed by Liu et al. (2019) and as part of the larger sample in van Velzen et al. (2020). Here, we provide analysis of the host galaxy in addition to a longer observational baseline with corresponding detailed analysis of the UV/optical photometric, optical spectroscopic, and X-ray properties of ASASSN-19dj. Throughout the paper, we assume a cosmology of $H_0 = 69.6$ km s $^{-1}$ Mpc $^{-1}$, $\Omega_M = 0.29$, and $\Omega_\Lambda = 0.71$. This paper is organized as follows. In Section 2, we detail the discovery and observations of the TDE. In Section 3, we

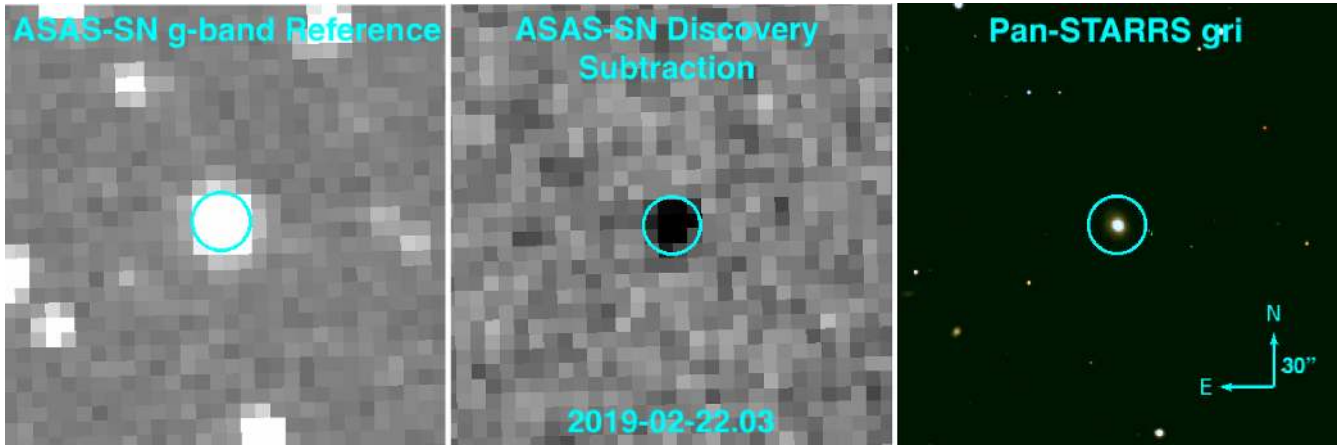


Figure 1. The ASAS-SN g -band reference image for the location of ASASSN-19dj (left), the subtracted ASAS-SN g -band discovery image from 2019 February 22.03 showing flux from ASASSN19-dj (centre), and a combined Pan-STARRS gri colour image of the host galaxy (right). The cyan circle (of radius 15.0 arcsec) marks the location of ASASSN-19dj.

present the analysis of our results. Section 4 provides a discussion of our results. Finally, our analysis is summarized in Section 5.

2 DISCOVERY AND OBSERVATIONS

ASASSN-19dj (α, δ) = (08:13:16.96, +22:38:54.00) was discovered in the g band in data from the ASAS-SN ‘Bohdan Paczyński’ unit in Cerro Tololo, Chile on 2019 February 22 (Brimacombe et al. 2019). Its discovery was announced on the Transient Name Server (TNS), and assigned the name AT 2019azh.¹ Rather than anonymise the discovering survey, in this paper, we will continue to refer to the TDE by its survey name ASASSN-19dj. ASASSN-19dj is located in the nucleus of the post-starburst galaxy KUG 0810+227, at a redshift of $z = 0.022346$ (Adelman-McCarthy et al. 2006). This redshift corresponds to a luminosity distance of 97.9 Mpc, making ASASSN-19dj one of the closest TDEs discovered to date. The g -band reference image used for host subtraction, the discovery image of ASASSN-19dj, and a false-colour Pan-STARRS gri image of the host galaxy² (Chambers et al. 2016) are shown in Fig. 1. The circle marking the location of ASASSN-19dj is 15.0 arcsec in radius, the same as the apertures used for the photometry presented in this paper.

Multiple spectroscopic observations were obtained shortly after discovery. Both the Nordic Optical Telescope Unbiased Transient Survey (NUTS; Heikkilä et al. 2019) and the extended Public ESO Spectroscopic Survey for Transient Objects (ePESSTO; Barbarino et al. 2019) obtained spectra that showed a strong blue continuum with few strong spectral features compared to the Sloan Digital Sky Survey (SDSS; York et al. 2000) host spectrum (see Section 2.1). The strong blue continuum, the appearance of broad $H\alpha$ emission lines, and a position consistent with the nucleus of the host galaxy made ASASSN-19dj a strong TDE candidate. Based on this, we triggered spectroscopic and ground-based photometric (Swope and LCOGT) follow-up of ASASSN-19dj.

Using ZTF and *Neil Gehrels Swift Gamma-ray Burst Mission* (*Swift*; Gehrels et al. 2004) observations, van Velzen et al. (2019) observed a plateau in the optical and UV light curve between 2019 February 24.25 and 2019 March 11.45. From fits to the ZTF and *Swift*

photometry, they found that the transient spectral energy distribution (SED) was consistent with a $T = (3.2 \pm 0.7) \times 10^4$ K blackbody, and measured a spatial separation from the host nucleus of 0.07 ± 0.31 arcsec. They classified the source as a TDE based on the observations of multiple blue spectra, a hot blackbody temperature, position in the centre of the host galaxy, and the lack of spectral features usually associated with AGN or supernovae. Using the central galaxy velocity dispersion from SDSS DR14 and the scaling relationship of Gültekin et al. (2009), they calculated an SMBH mass of $M_{\text{BH}} \gtrsim 4 \times 10^6 M_{\odot}$, and suggested that the observed plateau in the light curve was the result of Eddington-limited accretion.

2.1 Archival data of KUG 0810+227

KUG 0810+227 has been observed by several sky surveys across the electromagnetic spectrum. We obtained $ugriz$ and JHK_S images from SDSS Data Release 15 (Aguado et al. 2019) and the Two Micron All-Sky Survey (2MASS; Skrutskie et al. 2006), respectively. We measured aperture magnitudes using a 15.0 arcsec aperture radius in order to capture all of the galaxy light, and used several stars in the field to calibrate the magnitudes. We also obtained an archival *NUV* magnitude from the Galaxy Evolution Explorer (GALEX; Martin et al. 2005) All-sky Imaging Survey (AIS) catalogue and $W1$ and $W2$ magnitudes from the Wide-field Infrared Survey Explorer (WISE; Wright et al. 2010) AllWISE catalogue, giving us coverage from ultraviolet through mid-infrared wavelengths.

In order to constrain the possibility of the host galaxy being an AGN, we analysed a range of archival data for KUG 0810+227. Using *ROSAT* All-Sky Survey (RASS) data, we find no emission from the host galaxy at a 3σ upper limit of 3.4×10^{-2} counts s^{-1} . Assuming an AGN with a photon index of $\Gamma = 1.75$ (Ricci et al. 2017) and a Galactic column density of $N_{\text{H}} = 4.16 \times 10^{20} \text{ cm}^{-2}$ along the line of sight (HI4PI Collaboration 2016), this corresponds to an unabsorbed flux of $1.2 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ in the 0.3–10 keV band. At the distance of KUG 0810+227, this yields an X-ray luminosity of $1.4 \times 10^{42} \text{ erg s}^{-1}$. This limit rules out strong AGN activity, but does not rule out the presence of a weak or low-luminosity AGN (LLAGN; Tozzi et al. 2006; Marchesi et al. 2016; Liu et al. 2017; Ricci et al. 2017). The mid-infrared (MIR) colour of the host ($W1 - W2$) = 0.62 ± 0.04 mag again suggests that KUG 0810+227 does not harbour a strong AGN (e.g. Assef et al. 2013), but still does not

¹<https://wis-tns.weizmann.ac.il/object/2019azh>

²<http://ps1images.stsci.edu/cgi-bin/ps1cutouts?pos=123.320605325???2B2.648343&filter=colour>

Table 1. Archival host galaxy photometry.

Filter	Magnitude	Magnitude uncertainty
<i>NUV</i>	18.71	0.05
<i>u</i>	16.80	0.10
<i>g</i>	15.12	0.04
<i>r</i>	14.59	0.03
<i>i</i>	14.35	0.03
<i>z</i>	14.13	0.03
<i>J</i>	13.94	0.04
<i>H</i>	13.99	0.09
<i>K_S</i>	14.34	0.05
<i>W1</i>	15.07	0.03
<i>W2</i>	15.70	0.03

Note. Archival magnitudes of the host galaxy KUG 0810+227. *ugriz* and *JHK_S* magnitudes are 15.0 arcsec aperture magnitudes measured from SDSS and 2MASS images, respectively. The *NUV* magnitude is taken from the GALEX AIS and the *W1* and *W2* magnitudes are taken from the WISE AllWISE catalogue. All magnitudes are presented in the AB system.

rule out the presence of an LLAGN where the host light dominates over light from the AGN. When fitting a flat line to the WISE *W1* and *W2* light curves, we obtain reduced χ^2 values of 2.1 and 3.0, respectively, indicating a low level of variability consistent with an LLAGN.

We fit stellar population synthesis models to the archival photometry of KUG 0810+227 (shown in Table 1) using the Fitting and Assessment of Synthetic Templates (FAST; Kriek et al. 2009) to obtain an SED of the host. Our fit assumes a Cardelli, Clayton & Mathis (1988) extinction law with $R_V = 3.1$ and Galactic extinction of $A_V = 0.122$ mag (Schlafly & Finkbeiner 2011), a Salpeter IMF (Salpeter 1955), an exponentially declining star formation rate, and the Bruzual & Charlot (2003) stellar population models. Based on the FAST fit, KUG 0810+227 has a stellar mass of $M_* = 9.3^{+3.0}_{-1.2} \times 10^9 M_\odot$, an age of $1.4^{+0.6}_{-0.5}$ Gyr, and an upper limit on the star formation rate of $\text{SFR} \leq 6.9 \times 10^{-2} M_\odot \text{ yr}^{-1}$. The best-fitting age is slightly higher than the stellar ages of other TDE host galaxies (~ 0.5 Gyr; French, Arcavi & Zabludoff 2017). Using the sample of Mendel et al. (2014) to compute a scaling relation between stellar mass and bulge mass, we estimate a bulge mass of $\sim 10^{9.7} M_\odot$. We then use the $M_B - M_{\text{BH}}$ relation of McConnell & Ma (2013) to estimate a black hole mass of $\sim 10^{7.1} M_\odot$, roughly a factor of 3 higher than that estimated by van Velzen et al. (2019), although these methods use different data and scaling relations.

Our photometric follow-up campaign includes several filters for which archival imaging data are not available, including the *Swift* UVOT and *BV* filters. In order to estimate the host flux in these filters for host flux subtraction, we convolved the host SED from FAST with the filter response curve for each filter to obtain 15.0 fluxes. To estimate uncertainties on the estimated host galaxy fluxes, we perturbed the archival host fluxes assuming Gaussian errors and ran 1000 different FAST iterations. These synthetic fluxes were then used to subtract the host flux in our non-survey follow-up data.

The upper left panel of Fig. 2 compares the $H\alpha$ emission line equivalent width to the Lick $H\delta_A$ absorption index, which compares current and past star formation to identify post-starburst galaxies. The upper right panel of Fig. 2 shows the $H\alpha$ emission equivalent width as compared to the $\log_{10}([\text{N II}]/H\alpha)$ line ratio to separate ionization mechanisms, particularly those associated with LINER-like (Low-Ionization Nuclear Emission-line Region) emission line ratios. The bottom two panels of Fig. 2 show $\log_{10}([\text{O III}]/H\beta)$ versus $\log_{10}([\text{N II}]/H\alpha)$ and $\log_{10}([\text{O III}]/H\beta)$ versus $\log_{10}([\text{S II}]/H\alpha)$ line

ratios to characterize the activity of the host galaxies of TDEs. The background points in these figures are taken from the MPA-JHU catalogue (Brinchmann et al. 2004), which calculated the spectral properties of galaxies in SDSS DR8 (Eisenstein et al. 2011).

We obtained the archival SDSS (York et al. 2000) spectrum of KUG 0810+227. This spectrum shows $[\text{N II}] \lambda 6584$, $[\text{S II}] \lambda \lambda 6717, 6731$, $[\text{O I}] \lambda 6300$, and $[\text{O III}] \lambda \lambda 4959, 5007$ in emission with weak $H\alpha$ emission and $H\beta$ in absorption. We use the fits from the MPA-JHU catalogue (Brinchmann et al. 2004), which model and subtract the stellar component for robust emission line fits in our further analysis of KUG 0810+227. The line ratios $\log_{10}([\text{O III}]/H\beta) = 0.767$, $\log_{10}([\text{N II}]/H\alpha) = 0.004$, $\log_{10}([\text{S II}]/H\alpha) = -0.386$, and $\log_{10}([\text{O I}]/H\alpha) = -1.18$ place this galaxy in the AGN/Seyfert regions of the Baldwin et al. (1981) and Veilleux & Osterbrock (1987) diagrams ($[\text{O III}]/H\beta$ versus $[\text{N II}]/H\alpha$, $[\text{S II}]/H\alpha$, and $[\text{O I}]/H\alpha$). However, the WHAN diagram of Fig. 2 places KUG 0810+227 in the ‘retired galaxies’ (RG) region, where galaxies have ceased actively forming stars and are predominantly ionized by hot, evolved, lower mass stars such as post-AGB stars (Cid Fernandes et al. 2011). Other TDE hosts also tend to populate the RG region of this diagnostic diagram. Additionally, the observed line ratios of KUG 0810+227 can also be produced by large-scale shocks (e.g. Rich, Kewley & Dopita 2011, 2015). Thus, while KUG 0810+227 may harbour an LLAGN, it is also possible that other processes are at play.

From Fig. 2, we see that KUG 0810+227 is a post-starburst galaxy. The archival SDSS spectrum displays weak $H\alpha$ emission and extremely strong $H\delta$ absorption, with a Lick $H\delta_A$ index of $7.67 \pm 0.17 \text{ \AA}$ as measured by Brinchmann et al. (2004), confirming this classification. This is consistent with the tendency for TDEs to be found in post-starburst, or ‘quiescent Balmer-strong’, host galaxies (e.g. Arcavi et al. 2014; French et al. 2016; Law-Smith et al. 2017). Additionally, the host of ASASSN-19dj is similar to many other TDE hosts in terms of its star formation history. Compared to other TDE hosts, KUG 0810+227 is also similar in its line ratios. The possibility that KUG 0810+227 hosts an LLAGN is in line with the fact that hosts of other TDE such as ASASSN-14ae (Holoien et al. 2014a), ASASSN-14li (Holoien et al. 2016b; French et al. 2020), and ASASSN-19bt (Holoien et al. 2019b) show evidence for weak AGN activity.

The fact that KUG 0810+227 is an RG is in line with several studies on the ionization processes in post-starburst galaxies. De Propriis & Melnick (2014) compiled a sample of 10 post-starburst galaxies with *Hubble Space Telescope* (*HST*) imaging, optical spectra, X-ray, far-infrared, and radio data. They found no evidence of AGN down to an Eddington ratio of 0.1 per cent in these galaxies. Similarly, French et al. (2018) found that many post-starburst galaxies have LINER-like line ratios and that most are in the RG region of the WHAN diagram. The TDE hosts in particular have lower $H\alpha$ EW, placing them solidly in the RG region (French et al. 2017). However, Prieto et al. (2016) suggests that the host galaxy of the TDE ASASSN-14li, which we note is an RG in the WHAN diagram, may host an AGN.

Archival Catalina Real-Time Transient Survey (CRTS; Drake et al. 2009) data indicate that KUG 0810+227 experienced an outburst at MJD ~ 53640 (September 2005), roughly 14.5 yr prior to ASASSN-19dj. We obtained photometric data for this flare from both the Catalina Sky Survey (CSS) 0.7-m and the Mount Lemmon Survey (MLS) 1.5-m telescopes. First, we fit a flat line to the CRTS data between MJD = 54592 and MJD = 55919, outside of the flare to obtain a flux zero-point. The zero-point-subtracted light curve and a comparison of the flare to ASASSN-19dj are shown in Fig. 3. The reduced chi-squared of the non-outburst parts of the CRTS light curve as compared to the zero-point fit is 1.38, indicating that the

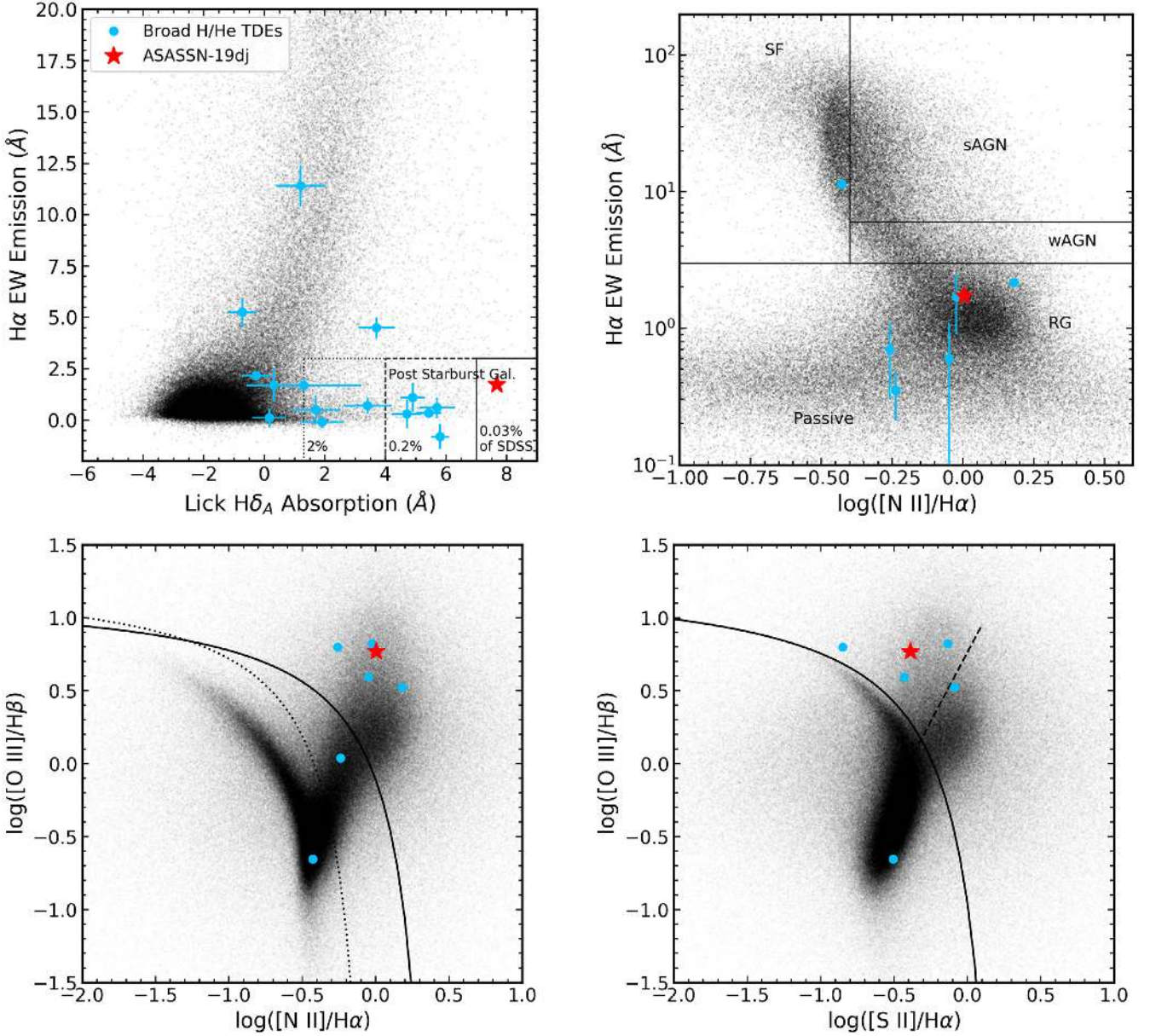


Figure 2. *Upper left panel:* H α emission line equivalent width (EW), which traces current star formation, as compared to the Lick H δ_A absorption index, which traces star formation in the past Gyr. The host galaxy KUG 0810+227 is shown as a red star, with other TDE hosts shown as blue circles. KUG 0810+227 is similar to some extreme post-starbursts galaxies seen in SDSS. The error bars on KUG 0810+227 are roughly the size of the symbol. *Upper right panel:* H α emission line equivalent width ($W_{H\alpha}$), as compared to $\log_{10}([N II]/H\alpha)$, otherwise known as the WHAN diagram (Cid Fernandes et al. 2011). Lines separating star-forming galaxies (SF), strong AGN (sAGN), weak AGN (wAGN), and passive and ‘retired galaxies’ (RG) are shown (Cid Fernandes et al. 2011). *Lower left panel:* $\log_{10}([O III]/H\beta)$ versus $\log_{10}([N II]/H\alpha)$ diagram (Baldwin, Phillips & Terlevich 1981; Veilleux & Osterbrock 1987). The solid line is the theoretical line separating AGNs (above right) and H II regions (below left) from Kewley et al. (2001). The dotted line is the empirical line from Kauffmann et al. (2003) showing the same separation. Objects between the dotted and solid lines are classified as composites. *Lower right panel:* $\log_{10}([O III]/H\beta)$ versus $\log_{10}([S II]/H\alpha)$ diagram (Veilleux & Osterbrock 1987). The solid line is the theoretical line separating AGN (above right) and H II regions (below left) from Kewley et al. (2001). The diagonal dashed line is the theoretical line separating Seyferts (above left) and LINERs (below right) from Kewley et al. (2006). KUG 0810+227 appears in the AGN/Seyfert region of both diagrams. In all panels, galaxies from SDSS Data Release 8 (Eisenstein et al. 2011) are shown in black.

V-band light curve is non-variable. The flare in CRTS appears to be of a similar magnitude to ASASSN-19dj, though the peak of the CRTS flare may have occurred in the seasonal gap. The FWHM (full width at half-maximum) of the ASASSN-19dj flare is ~ 80 d in the ASAS-SN g -band data, while the FWHM of the CRTS flare is $\lesssim 140$ d, although we do not see the full rise or peak of this flare. Finally, the light curve of ASASSN-19dj appears to decline slower than the CRTS flare as indicated by the right-hand panel of Fig. 3.

In addition to our search of CRTS data, we searched archival Pan-STARRS data for previous outbursts. In this search, we found no clear detections prior to ASASSN-19dj. This places strong constraints on any potential flaring activity of the host galaxy, with two large (~ 15 – 16 mag) flares and otherwise no significant optical variability over the course of roughly 15 yr.

The photometry from CRTS uses SExtractor (Bertin & Arnouts 1996), which includes flux from the entire host galaxy. Thus it is not

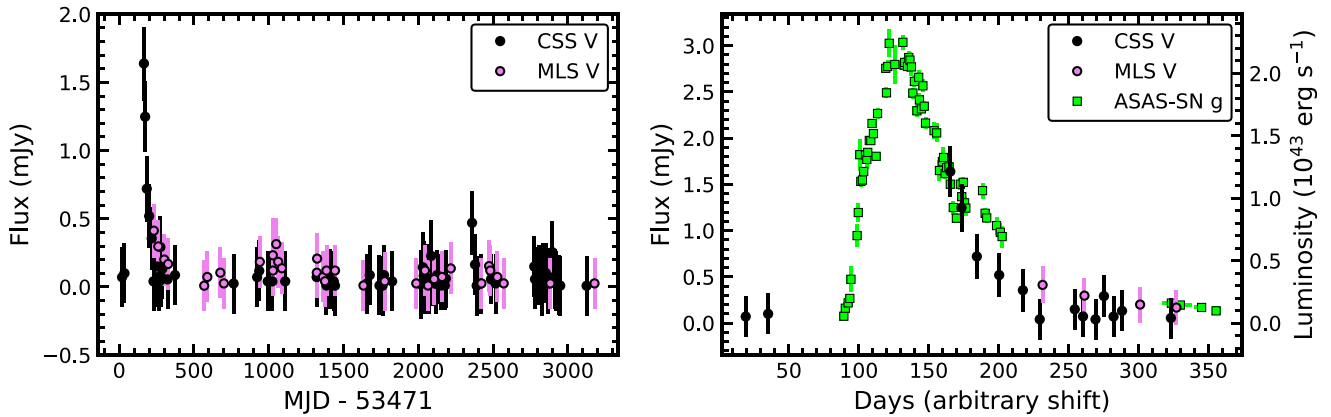


Figure 3. *Left-hand panel:* The V-band CRTS data (CSS in black and MLS in violet) light curves. *Right-hand panel:* Comparison of the flare seen at MJD \sim 53680 by CRTS and the ASAS-SN *g*-band data (shown in green) of the TDE ASASSN-19dj with an arbitrary shift in days between the two flares so that they are roughly aligned. The two light curves are offset by their best-fitting zero-points and have been corrected for Galactic extinction.

possible to determine the nature of the transient from the CRTS photometry alone. In an effort to link this flare to a known transient, we searched multiple data bases of known supernovae such as the Open Supernova Project (Guillochon et al. 2017b) and the Central Bureau for Astronomical Telegrams (CBAT³) but found no reported SNe or other flares at the time. Nevertheless, even if the source is consistent with the nucleus, it is difficult to constrain the nature of the earlier transient without spectroscopic or multiband photometric observations during that epoch.

2.2 ASAS-SN light curve

ASAS-SN is a fully automated transient survey which consists of 20 telescopes on five robotic mounts. Each telescope is a 14-cm aperture Nikon telephoto lens with 8.0 arcsec pixels, and each unit consists of four telescopes on a common mount. Single ASAS-SN units are located at Haleakalā Observatory, McDonald Observatory, and the South African Astrophysical Observatory (SAAO), and two are located at Cerro Tololo Inter-American Observatory (CTIO). With our current network, ASAS-SN monitors the visible sky with a cadence of \sim 20 h to a depth of $g \sim 18.5$ mag.

Since ASASSN-19dj is near the equator, it was observed from both hemispheres and by all five units. Additionally, ASASSN-19dj lies in a designed 0.5 deg field overlap region, giving us roughly twice the cadence. Thus we observed ASASSN-19dj with 10 of the 20 ASAS-SN cameras currently deployed. However, the filters for two of the ASAS-SN units, Brutus and Cassius, had been changed from *V* to *g* shortly before ASASSN-19dj was discovered. Thus, there were not enough *g*-band images taken before the TDE to construct a subtraction reference image for those cameras. Because of this, we had to modify the normal ASAS-SN processing pipeline to extract the light curve.

Images were reduced using the automated ASAS-SN pipeline but we performed image subtraction separately. We used the ISIS image subtraction package (Alard & Lupton 1998; Alard 2000) with the same parameters as the ASAS-SN pipeline, but images from all cameras for a given pointing were first interpolated on to a common grid. We then built a reference image using good images from

multiple cameras observed well before the rise of ASASSN-19dj. This common reference image was used to analyse all the data.

We then used the IRAF APPHOT package with a 2-pixel radius (approximately 16.0 arcsec) aperture to perform aperture photometry on each subtracted image, generating a differential light curve. The photometry was calibrated using the AAVSO Photometric All-Sky Survey (Henden et al. 2015). We visually inspected each of the 1155 exposures taken after 2018 February 5 analysed in this work for clouds or flat-fielding issues and disregarded images where issues were seen. We also discarded images with an FWHM of 1.67 pixels or greater.

To increase the signal-to-noise ratio (S/N), we stacked our photometric measurements. For exposures with TDE emission, the photometric measurements were stacked in 12 hour bins. Prior to our first detection, measurements within 75 hours of each other were stacked to get deeper upper limits. After the 2019 seasonal gap, measurements within 10 d of each other were stacked to sample the decline of the TDE.

2.3 ATLAS light curve

The ATLAS survey is designed primarily to detect small (10–140 m) asteroids that may collide with Earth (Tonry et al. 2018). ATLAS uses two 0.5 m *f*/2 Wright Schmidt telescopes on Haleakalā and at the Mauna Loa Observatory. For normal operation, the telescopes obtain four 30-second exposures of 200–250 fields per night. This allows the telescopes to cover roughly a quarter of the sky visible from Hawaii each night, ideal for transient detection (Smith et al. 2020). ATLAS uses two broad-band filters, the ‘cyan’ (*c*) filter from 420–650 nm and the ‘orange’ (*o*) filter covering the 560–820 nm range (Tonry et al. 2018).

Each ATLAS image is processed by a pipeline that performs flat-field corrections in addition to astrometric and photometric calibrations. Reference images of the host galaxy were created by stacking multiple images taken under excellent conditions before MJD = 58251 and this reference was then subtracted from each science image of ASASSN-19dj in order to isolate the flux from the transient. We performed forced photometry on the subtracted ATLAS images of ASASSN-19dj as described in Tonry et al. (2018). We combined the four intra-night photometric observations using a weighted average to get a single flux measurement. The ATLAS

³<http://www.cbat.eps.harvard.edu/lists/Supernovae.html>

Table 2. Host-subtracted photometry of ASASSN-19dj.

MJD	Filter	Magnitude	Uncertainty	Telescope
58537.07	<i>i</i>	16.36	0.01	LCOGT-1m
58539.86	<i>i</i>	16.24	0.01	LCOGT-1m
58541.81	<i>i</i>	16.12	0.01	LCOGT-1m
...
58850.21	<i>UVW2</i>	17.68	0.07	<i>Swift</i>
58906.65	<i>UVW2</i>	17.74	0.11	<i>Swift</i>
58911.89	<i>UVW2</i>	18.00	0.18	<i>Swift</i>

Note. Host-subtracted magnitudes and 3σ upper limits for all follow-up photometry. A range of MJD in the first column indicates the beginning and end of the range over which data were stacked to increase S/N. All magnitudes are corrected for Galactic extinction and presented in the AB system. The last column reports the source of the data for each epoch. The *Swift B* data do not include the shift applied in Fig. 4. Only a small section of the table is displayed here. The full table can be found online as an ancillary file.

o-band photometry and 3σ limits are presented in Table 2 and are shown in Fig. 4. We do not plot the *c*-band photometry in Fig. 4 as there were few *c*-band observations in the rise to peak and near peak due to weather and the design of the ATLAS survey, but as they provide useful early limits and detections, we present them in Table 2.

2.4 ZTF light curves

The ZTF survey uses the Samuel Oschin 48-in Schmidt telescope at Palomar Observatory and a camera with a 47 deg^2 field of view that reaches as deep as 20.5 *r*-band mag in a 30 s exposure. Alerts for transient detection from ZTF are created from the final difference images (Patterson et al. 2019). These alerts are distributed to brokers including Lasair (Smith et al. 2019) through the University of Washington Kafka system. For ASASSN-19dj, we obtained ZTF *g*- and *r*-band light curves from the Lasair broker.⁴ Lasair uses ZTF difference imaging photometry so the host flux is subtracted. The ZTF magnitudes presented in this paper are calculated using PSF photometry. Similar to the ATLAS data, we combined the intra-night photometric observations using a weighted average to get a single flux measurement.

2.5 Additional ground-based photometry

We also obtained photometric follow-up observations from several ground-based observatories. We used the Las Cumbres Observatory (Brown et al. 2013) 1-m telescopes located at CTIO, SAAO, McDonald Observatory, and Siding Spring Observatory for *BVgri* observations, and the Swope 1-m telescope at Las Campanas Observatory for *uBVgri* observations. After applying flat-field corrections, we solved astrometry in each image using astrometry.net (Barron et al. 2008; Lang et al. 2010).

We aligned the *ugri* data to the archival SDSS image in the corresponding filter for each follow-up image using the Python REPROJECT package, which uses the WCS information of two images to project one image on to the other. We then subtracted the SDSS template images from each follow-up image using HOTPANTS⁵ (Becker 2015), an implementation of the Alard (2000) image subtraction algorithm, and used the IRAF APPHOT package to measure 5.0 arcsec aperture

magnitudes of the transient. For the *BV* data, we did not have archival images available to use as subtraction template images. Instead, we used APPHOT to measure 15.0 arcsec aperture magnitudes of the host + transient, and subtracted the 15.0 arcsec host flux synthesized from our FAST fit in the appropriate filter to isolate the transient flux. For all filters, we used SDSS stars in the field to calibrate our photometry, using the corrections from Lupton (2005) to calibrate the *B* and *V* band magnitudes with the *ugriz* data.

We measured the centroid position of the transient in a host-subtracted Las Cumbres Observatory *g*-band image taken near peak using the IRAF IMCENTROID package. This yielded a position of $(\alpha, \delta) = (08:13:16.96, +22:38:54.00)$. We also used the archival SDSS *g*-band image to measure the position of the nucleus of KUG 0810+227, finding $(\alpha, \delta) = (08:13:16.95, +22:38:53.89)$. This gives an angular offset of 0.21 ± 0.12 arcsec, where the uncertainty is due to uncertainty in the centroid positions of the TDE and host nucleus. We also measured the centroid positions of several stars in both the follow-up and SDSS host images, finding that the stars had an average random offset of 0.19 arcsec. Combining these sources of uncertainty, the transient position is thus offset 0.21 ± 0.24 arcsec from the position of the host nucleus, corresponding to a physical distance of 99.2 ± 112.5 pc.

2.6 Swift observations

Forty-four total Neil Gehrels Swift Gamma-ray Burst Mission (*Swift*; Gehrels et al. 2004) target-of-opportunity (ToO) observations were carried out between 2019 March 2 and 2020 March 3 Swift target ID 11186 (as AT2019azh; PIs: Gezari, Arcavi, and Wevers), and Swift target ID 12174 (as ASASSN-19dj; PI: Hinkle). These observations used the UltraViolet and Optical Telescope (UVOT; Roming et al. 2005) and X-ray Telescope (XRT; Burrows et al. 2005) to study the multiwavelength properties of the TDE.

2.6.1 UVOT observations

For a majority of the observation epochs, *Swift* observed ASASSN-19dj with all six UVOT filters (Poole et al. 2008): *V* (5468 Å), *B* (4392 Å), *U* (3465 Å), *UVW1* (2600 Å), *UVM2* (2246 Å), and *UVW2* (1928 Å). Each epoch of UVOT data includes two observations in each filter, which we combined into one image for each filter using the HEASOFT UVOTIMSUM package. We then used the UVOTSOURCE package to extract source counts using a 15.0 arcsec radius region centred on the position of the TDE and background counts using a source-free region with radius of ~ 40.0 arcsec. We converted the UVOT count rates into fluxes and magnitudes using the most recent calibrations (Poole et al. 2008; Breeveld et al. 2010).

Because the UVOT uses unique *B* and *V* filters, we used publicly available colour corrections⁶ to convert the UVOT *BV* data to the Johnson–Cousins system. We then corrected the UVOT photometry for Galactic extinction and removed host contamination by subtracting the corresponding 15.0 arcsec host flux in each filter, as we did with the ground-based *BV* data.

Fig. 4 shows the extinction-corrected, host-subtracted light curves of ASASSN-19dj. The photometry spans from the shortest *UVW2* (1928 Å) band of *Swift* to *i* band (7609 Å) from Swope and LCOGT and includes the data ranging from 21 d prior to peak to 392 d after

⁴<https://lasair.roe.ac.uk/>

⁵<http://www.astro.washington.edu/users/becker/v2.0/hotpants.html>

⁶https://heasarc.gsfc.nasa.gov/docs/heasarc/caldb/swift/docs/uvot/uvot_cal_db_coltrans_02b.pdf

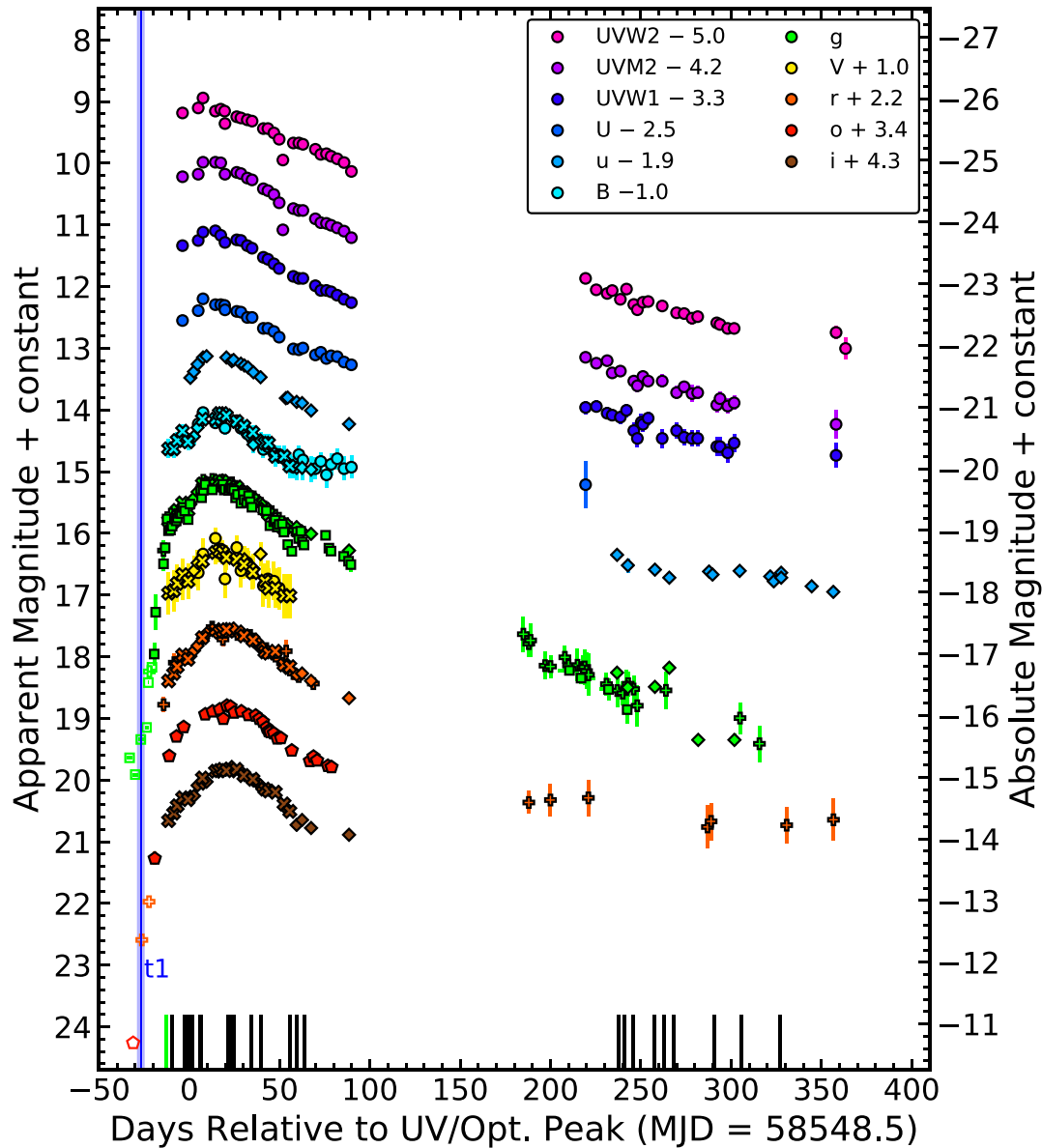


Figure 4. Host-subtracted UV and optical light curves of ASASSN-19dj, showing ASAS-SN (*g*, squares), Swift (UV+UBV, circles), ATLAS (*o*, pentagons), ZTF (*gr*, pluses), Swope (*uBVgri*, diamonds), and Las Cumbres Observatory 1-m telescopes (*BVgri*, x-shapes) photometry spanning from roughly 20 d prior to peak (MJD = 58548.5) to roughly 390 d after in observer-frame days. Horizontal error bars on the ASAS-SN *g*-band data indicate the date range of observations stacked to obtain deeper limits and higher S/N detections, although they are small and difficult to see. Open symbols indicate 3σ upper limits when the TDE was not detected. The green bar on the *x*-axis marks the epoch of ASAS-SN discovery. Black bars along the *x*-axis show epochs of spectroscopic follow-up. The blue line is the estimated time of first light (see Section 3.1) with the shading corresponding to the uncertainty. All data are corrected for Galactic extinction and shown in the AB system.

peak. The corrected *Swift B* data were inconsistent with the ground-based *B* data, so we shifted the *Swift B* data in Fig. 4. To do this, we computed the median offset between the LCOGT and Swope *B*-band data from the *Swift B* data, which was found to be 0.27 mag. All the UV and optical photometry shown in Fig. 4, in addition to limits not shown in this figure, is presented in Table 2.

2.6.2 XRT observations

In addition to the *Swift* UVOT observations, we also obtained simultaneous *Swift* X-Ray Telescope (XRT) photon-counting observations.

All observations were reprocessed from level one XRT data using the *Swift* XRTPipeline version 0.13.2, producing cleaned event files and exposure maps. Standard filter and screening criteria⁷ were used, as well as the most up-to-date calibration files.

To extract both background-subtracted count rates and spectra, we used a source region with a radius of 50 arcsec centred on the position of ASASSN-19dj and a source free background region centred at $(\alpha, \delta) = (08^{\text{h}}13^{\text{m}}07^{\text{s}}.93, +22^{\circ}35'15''.36)$ with a radius of 150.0 arcsec. The reported count rates are aperture corrected where a

⁷http://swift.gsfc.nasa.gov/analysis/xrt_swguide_v1_2.pdf

50 arcsec source radius contains ~ 90 per cent of the source counts at 1.5 keV, assuming an on-axis pointing (Moretti et al. 2004). To increase the S/N of our observations, we combined our individual *Swift* observations into six time bins using XSELECT version 2.4g, allowing us to extract spectra with $\gtrsim 200$ –300 background-subtracted counts.

To extract spectra from our merged observations, we used the task XRTPRODUCTS version 0.4.2 and the regions defined above to extract both source and background spectra. To extract ancillary response files (ARFs), we first merged the corresponding individual exposure maps that were generated by XRTPIPELINE using XIMAGE version 4.5.1 and then used the task XRTMKARF. We used the ready-made response matrix files (RMFs) that are available with the *Swift* calibration files. Each spectrum was grouped to have a minimum of 10 counts per energy bin using the FTOOLS command GRPPHA.

2.7 NICER observations

After the 2019 seasonal gap, follow-up *Swift* XRT observations found that the X-ray flux of ASASSN-19dj had increased by an order of magnitude compared to the flux approximately 100 d earlier (see Section 3.5). ToO observations of ASASSN-19dj were then obtained using the Neutron star Interior Composition Explorer (NICER; Gendreau, Arzoumanian & Okajima 2012), which is an external payload on the International Space station that has a large effective area over the 0.2–12.0 keV energy band and provides fast X-ray timing and spectroscopic observations of sources. In total, 80 observations were taken between 2019 October 23 and 2020 March 12 (Observation IDs: 2200920101–2200920176, 3200920101–3200920105, PI:Pasham/Gendreau, Pasham et al. 2019), totaling 169 ks of cumulative exposure.

The data were reduced using NICERDAS version 6a, HEASOFT version 6.26.1. Standard filtering criteria were applied using the NICERDAS task NICERL2. Here the standard filter criteria includes⁸: the *NICER* pointing is (ANG_DIST) < 0.015 degrees from the position of the source; excluding events that were acquired during passage through the South Atlantic Anomaly, or those that are obtained when Earth was 30° (40°) above the dark (bright) limb (ELV and BR.EARTH, respectively). We also removed events that are flagged as overshoot, or undershoot events (EVENT_FLAGS = bxxxx00), and we used the so-called ‘trumpet filter’ to remove events with a PI_RATIO $> 1.1 + 120/\text{PI}$, where PI is the pulse invariant amplitude of an event, as these are likely particle events (Bogdanov et al. 2019). To extract time-averaged spectra and count rates, we used XSELECT, and readymade ARF (nixtiaveonaxis20180601v002.arf) and RMF (nixtref20170601v001.rmf) files that are available with the *NICER* CALDB. Similar to the *Swift* spectra, each spectrum was grouped with a minimum of 10 counts per energy bin. As *NICER* is a non-imaging instrument, background spectra were generated using the background modelling tool NIBACKGEN3C50.⁹

To analyse the spectra extracted from both our *Swift* and *NICER* observations, we used the X-ray spectral fitting package (XSPEC) version 12.10.1f (Arnaud 1996) and χ^2 statistics. Both the *Swift* and *NICER* data and their analysis are further discussed in Section 3.5.

⁸See https://heasarc.gsfc.nasa.gov/docs/nicer/data_analysis/nicer_analysis_guide.html or (Bogdanov et al. 2019) for more details about these criteria.

⁹https://heasarc.gsfc.nasa.gov/docs/nicer/tools/nicer_bkg_est_tools.html

Table 3. X-ray luminosity and hardness ratios of ASASSN-19dj.

MJD	log Lum. (erg s ⁻¹)	Lum. error (erg s ⁻¹)	HR	HR error	Satellite
58544.76	40.72	–	–0.07	–	<i>Swift</i>
58553.45	41.42	0.23	–0.53	0.28	<i>Swift</i>
58556.11	41.36	0.26	–1.00	0.01	<i>Swift</i>
...
58922.58	42.06	0.04	–0.70	0.09	<i>NICER</i>
58934.02	42.03	0.06	–0.63	0.11	<i>NICER</i>
58940.53	41.96	0.06	–0.40	0.07	<i>NICER</i>

Note. X-ray luminosities and hardness ratios with associated uncertainties. Dashed lines represent 3σ upper limits. The hardness ratio is defined as $(H - S)/(H + S)$, where we define hard counts H as the number of counts in the 2–10 keV range and soft counts S are the number of counts in the 0.3–2 keV. The last column reports the source of the data for each epoch. Only a small section of the table is displayed here. The full table can be found online as an ancillary file.

2.8 XMM–Newton slew observations

In addition to *Swift* and *NICER* observations, we also searched for *XMM–Newton* slew observations that overlap the position of ASASSN-19dj. These slew observations are taken using the PN detector of *XMM–Newton* as it manoeuvres between pointed observations, detecting X-ray emission down to a 0.2–10.0 keV flux limit of $\sim 10^{-12}$ erg cm⁻² s⁻¹ (Saxton et al. 2008). We found two slew observation (ObsIDs: 9353900003 and 9363000003) coincident with the position of ASASSN-19dj. These observations were taken on 2019 April 7 and 2019 October 5, respectively, corresponding to ~ 44 and ~ 225 d after discovery. To analyse these observations, we follow the current slew analysis thread on the *XMM–Newton* Science System (SAS) data analysis threads.¹⁰ Here, we use the SAS tool command ESLEWCHAIN and the most up to date calibration files to produce filtered event files that we use in our analysis. Similar to our *XRT* analysis, we extract the number of counts using a source region with a radius of 30 arcsec centred on the position of ASASSN-19dj and a source free background region with a radius of 150 arcsec centred at $(\alpha, \delta) = (08^{\text{h}} 13^{\text{m}} 31^{\text{s}}.79, +22^{\circ} 37' 30''.53)$. A 30 arcsec source region contains 85 per cent of all source photons at 1.9 keV. Due to the low exposure times of each observation, which was determined using the corresponding exposure files of each observation, no spectra could be extracted. The 0.3–10 keV X-ray luminosities and hardness ratios derived from the count rates for the various X-ray epochs are shown in Table 3.

2.9 Spectroscopic observations

In addition to the ePESSTO spectrum released on TNS, we obtained follow-up spectra of ASASSN-19dj with the Low Dispersion Survey Spectrograph (LDSS-3) on the 6.5-m Magellan Clay telescope, the Inamori-Magellan Areal Camera and Spectrograph (IMACS; Dressler et al. 2011) on the 6.5-m Magellan-Baade telescope, the Wide Field Reimaging CCD Camera (WFCCD) on the du Pont 100-inch telescope, the SuperNova Integral Field Spectrograph (SNIFS; Lantz et al. 2004) on the 88-in University of Hawaii telescope, the Low-Resolution Imaging Spectrometer (LRIS; Oke et al. 1995) on the 10-m Keck I telescope, the Goodman High Throughput Spectrograph (GHTS; Clemens, Crain & Anderson 2004) on the

¹⁰<https://www.cosmos.esa.int/web/xmm-newton/sas-thread-epic-slew-processing>

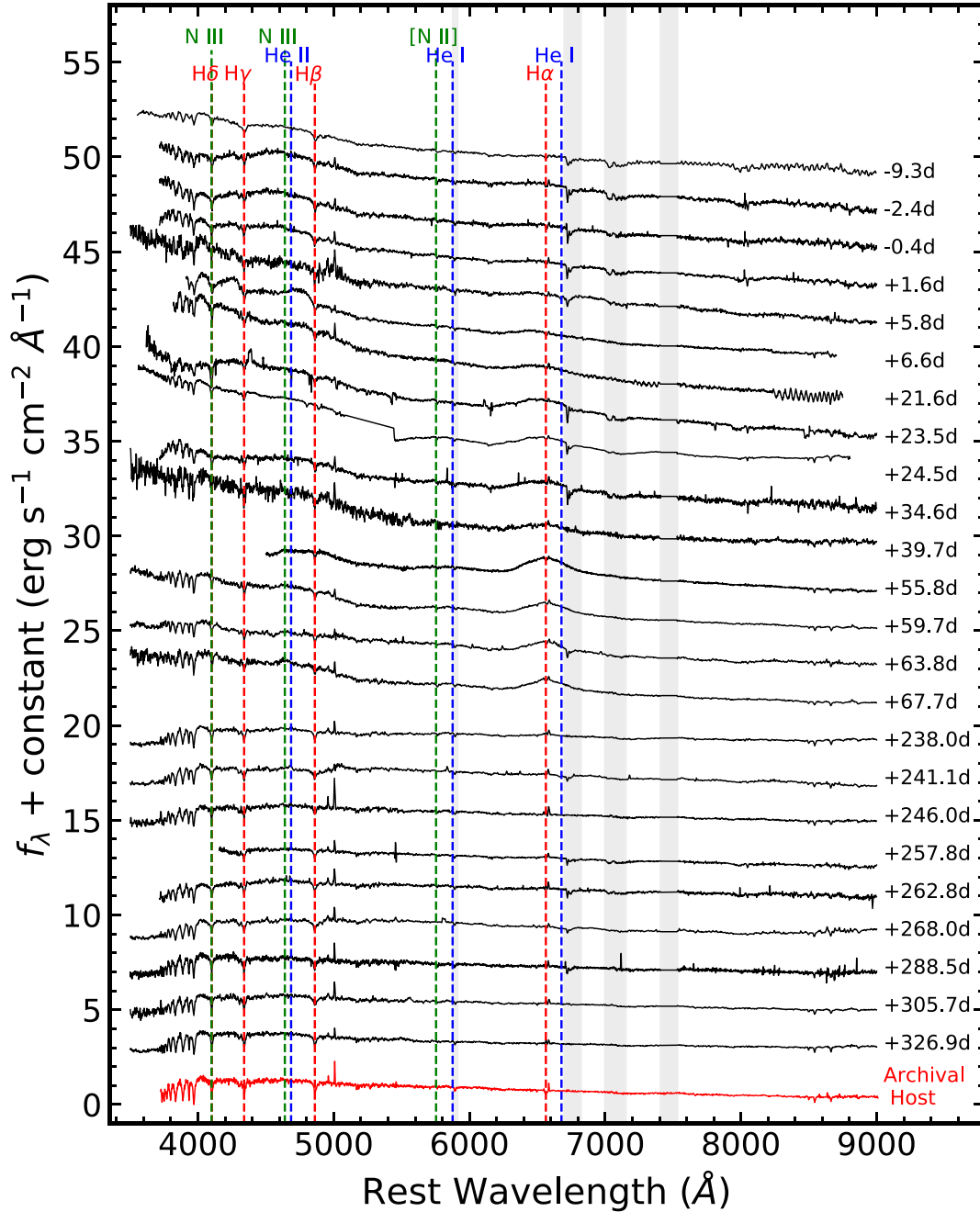


Figure 5. Optical spectroscopic evolution of ASASSN-19dj spanning from 9.3 d prior to peak UV/optical emission (top) until 326.9 d after peak (bottom). These spectra are calibrated using the photometry presented in Fig. 4. The vertical grey bands mark atmospheric telluric features and the strong telluric feature between ~ 7400 and 7550 Å has been masked. An archival host spectrum from SDSS is shown in red at the very bottom. The vertical lines mark spectral features common in TDEs, with hydrogen lines in red, helium lines in blue, and nitrogen lines in green. The straight line in the spectrum at 24.5 d after peak connects the blue and red sides of the LRIS spectrum with a large dichroic feature.

4.1-m Southern Astrophysical Research Telescope (SOAR), the Kast Double spectrograph on the Lick Shane 120-in telescope, and the Multi-Object Double Spectrographs (MODS; Pogge et al. 2010) on the Large Binocular Telescope (LBT; Hill, Green & Slagle 2006). Three of our spectra were obtained prior to peak light and 21 were obtained after peak. Most of the spectra were reduced and calibrated with standard IRAF procedures, such as bias subtraction, flat-fielding, 1D spectroscopic extraction, and wavelength calibration.

The IMACS data from 2019 November 19.3 were reduced using an updated version of the routines developed by Kelson et al. (2014). The flux calibration for our observations was initially done using standard star spectra obtained on the same nights as the science spectra. We then scaled and tilted our spectra to match the calibrated flux of the TDE in the optical photometry.

All the classification and follow-up spectra for ASASSN-19dj are presented in Fig. 5. From top to bottom, the optical spectrum

evolves from a hot, blue continuum to a host-dominated spectrum. The locations of several emission lines commonly seen in TDEs are marked with vertical dashed lines. Some of these emission lines appear, evolve, and disappear throughout the time period probed by these spectra.

3 ANALYSIS

3.1 Light curve

Using Markov chain Monte Carlo (MCMC) methods, we fit each of the epochs where there is *Swift* UV photometry as a blackbody to obtain the bolometric luminosity, temperature, and effective radius of ASASSN-19dj. So that our fits are relatively unconstrained, we ran each of our blackbody fits with flat temperature priors of $10\,000\text{ K} \leq T \leq 55\,000\text{ K}$. To find the time of peak UV/optical luminosity, we fit a parabola to the light curve created by bolometrically correcting the ASAS-SN *g*-band light curve using these blackbody fits. For this fit, we excluded any upper limits. Because the curve is quite flat near peak, we fit the parabola in a narrow range between $\text{MJD} = 58535.2$ and $\text{MJD} = 58556.2$. We generated 10 000 realizations of the bolometric light curve in this date range with each magnitude perturbed by its uncertainty assuming Gaussian errors. We then fit a parabola to each of these light curves and took the median value as the peak and 16th and 84th percentiles as the uncertainties in peak time. Using this procedure, we find the time of peak bolometric luminosity to be $\text{MJD} = 58548.5^{+6.3}_{-2.6}$. From Fig. 4, looking from the shortest wavelength (*UVW2*) to the longest (*i*), we see that the time of peak light in each band is offset. Using a similar procedure to the bolometric light curve, but for the flux in a single photometric band, we find that the *Swift UVW2* light-curve peaks at $\text{MJD} = 58554.9^{+1.1}_{-1.5}$ and the LCOGT *i*-band light curve peaks at $\text{MJD} = 58571.9 \pm 0.1$. This offset of ~ 17 d is longer than the offsets seen in other TDEs such as ASASSN-18pg (Holoien et al. 2020) and ASASSN-19bt (Holoien et al. 2019b). This likely occurs because the temperature of ASASSN-19dj steadily declines for roughly 25 d before increasing for 20 d and gradually levelling off for the next ~ 230 d (see Section 3.3), in contrast to the TDEs that exhibit relatively constant temperature near peak (e.g. Hinkle et al. 2020; van Velzen et al. 2020).

ASASSN-19dj is one of only a few TDEs for which the early-time coverage is adequate to fit a rise slope. We fit the early-time rise as a power law with

$$f = z \text{ for } t < t_1, \text{ and} \quad (1)$$

$$f = z + h \left(\frac{t - t_1}{\text{days}} \right)^\alpha \text{ for } t > t_1. \quad (2)$$

This model fits for the zero-point z , the time of first-light t_1 , a flux scale h , and the power-law index α . An MCMC fit yields the best-fitting parameters $z = -30.7^{+8.6}_{-9.2} \mu\text{Jy}$, $h = 7.4^{+14.0}_{-5.5} \mu\text{Jy}$, $t_1(\text{MJD}) = 58521.9^{+1.3}_{-1.7}$, and $\alpha = 1.90^{+0.42}_{-0.36}$. These fits are shown in Fig. 6.

ASASSN-19dj is only the third TDE for which a power-law could be fit to the early-time light curve. This best-fitting power-law index of $\alpha = 1.90^{+0.42}_{-0.36}$ is consistent with the fireball model used for the early-time evolution of SNe (e.g. Riess et al. 1999; Nugent et al. 2011). ASASSN-19bt, the TDE with the best early-time data also has a rise consistent with this model (Holoien et al. 2019b). Unlike the model invoked for SNe, where the ejecta initially expands at a constant velocity and temperature, the early stages of a TDE are more complex, so it is somewhat odd that these two objects have shown such a rise. Further analysis of more early-time TDE light curves will help us better understand their rise slopes.

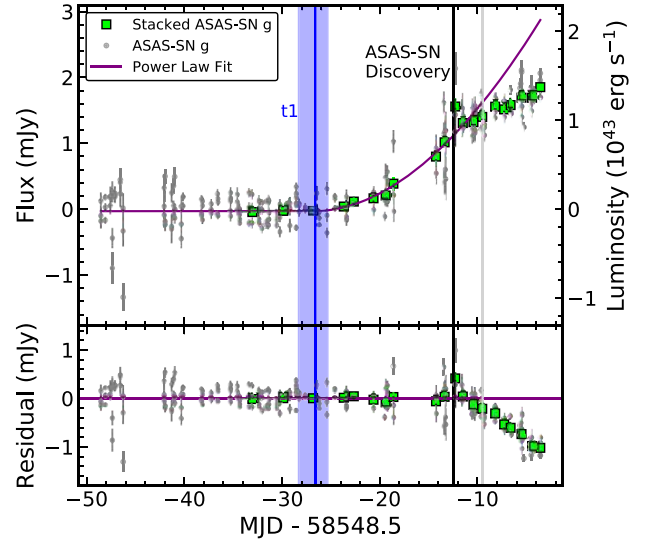


Figure 6. *Top panel:* Stacked (green) and raw (grey) ASAS-SN *g*-band light curve and best-fitting power-law model in purple. This power-law fit yields a time of first light of $t_1 = 58521.9^{+1.3}_{-1.7}$ and a power-law index of $\alpha = 1.90^{+0.42}_{-0.36}$. The blue line shows the fitted time of first light with the shading representing the uncertainty and the black line shows the ASAS-SN discovery date. The light grey line marks the last epoch that was fit by our MCMC model. *Bottom panel:* Residuals between the data and best-fitting power-law model. The ASAS-SN *g*-band light curves deviates from a power-law rise roughly 16 d after the fitted time of first light.

Additionally, the rise slope of the early-time bolometric light curve of ZTF19abzrhgq/AT2019qiz was found to be similar to ASASSN-19dj and ASASSN-19bt (Nicholl et al. 2020). It has been suggested that a t^2 rise may be the result of an outflow (e.g. Nicholl et al. 2020). The existence of blue-shifted emission lines early in the evolution of ASASSN-19dj may support this possibility. Finally, the ASAS-SN *g*-band light curve follows a t^2 power-law rise for approximately 16 d, similar to ASASSN-19bt (Holoien et al. 2019b).

From Fig. 6, we see that the light curve rises from the time of first light to the peak UV/optical bolometric luminosity in 26 d, shorter than the rise to peak time measured for ASASSN-19bt (Holoien et al. 2019b) and the limits on rise times for PS18kh (Holoien et al. 2019a), and ASASSN-18pg (Holoien et al. 2020). Assuming that the early-time X-rays are due to accretion and are evidence of prompt circularization (see Section 3.5), this more rapid rise may indicate a more efficient circularization of material for ASASSN-19dj than other TDEs. From the fitted time of first light, we find that ASAS-SN discovered this transient within about 2 weeks of the beginning of the flare.

We used the Modular Open-Source Fitter for Transients (MOSFiT; Guillochon et al. 2017a; Mockler et al. 2019) to fit the host-subtracted light curves of ASASSN-19dj to estimate physical parameters of the star, SMBH, and the encounter. MOSFiT uses models containing several physical parameters to generate bolometric light curves of a transient, generates single-filter light curves from the bolometric light curves, and fits these to the observed multiband data. It then finds the combination of parameters yielding the highest likelihood match for a given model using one of various sampling methods. We ran the MOSFiT TDE model in nested sampling mode when fitting our data, as we have a large number of observations in several photometric filters.

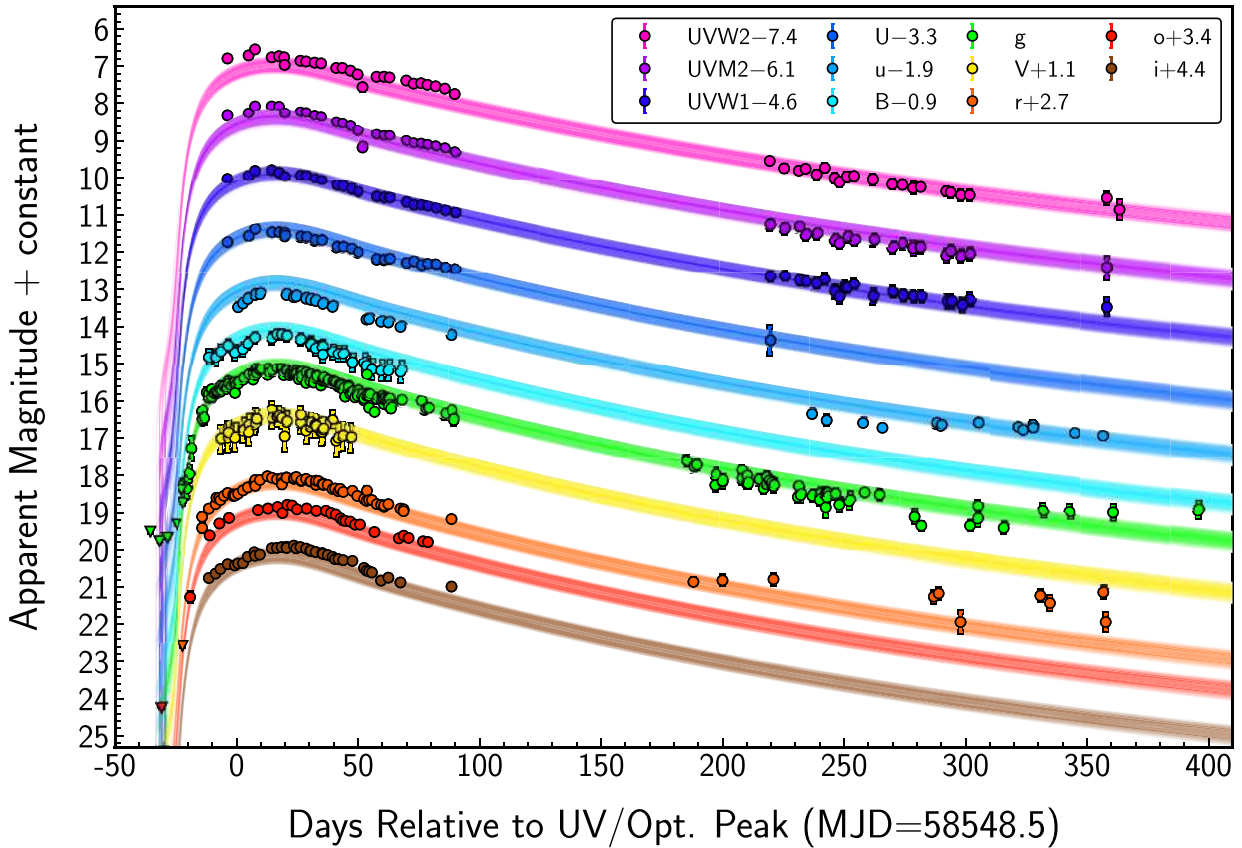


Figure 7. MOSFiT light-curve fits and host-subtracted light curves. The 1–99 per cent range of fits for each filter are shown as shaded regions with the median fit shown as a solid line. All detections are plotted as circles with 3σ upper limits plotted as downward triangles. The colours match those of Fig. 4.

MOSFiT is the only available tool for generalized fitting of TDE emission, and seems to model cases such as ASASSN-19dj, which has relatively smooth light curves, fairly well. We show the MOSFiT multiband fits to the ASASSN-19dj light curves in Fig. 7 with our data overplotted. MOSFiT does a reasonable job of fitting the data near the peak, particularly in the optical, though it slightly underestimates the UV emission near peak. The rise is particularly well constrained compared to many of the TDEs in the sample fit by Mockler et al. (2019). The fit to the late-time UV data is good, but many of the optical bands appear to flatten relative to the MOSFiT decline. The UV excess near peak and the late-time optical flattening may indicate the presence of multiple emission components throughout the evolution of ASASSN-19dj.

Table 4 shows the median values and 1–99 per cent range for the MOSFiT TDE model parameters. The model parameters are generally very well constrained, with statistical uncertainties from the fit being significantly smaller than the systematic uncertainties of the model (see table 3 of Mockler et al. 2019). The black hole mass and stellar mass given by MOSFiT are $M_h = 7.8^{+3.9}_{-4.1} \times 10^6 M_\odot$ and $M_* = 0.10^{+0.37}_{-0.08} M_\odot$, respectively. This black hole mass is larger than, but marginally consistent with, the mass limit calculated by van Velzen et al. (2019) and consistent with our estimate in Section 2.1. The stellar mass, while low, is consistent with several other TDEs modelled in Mockler et al. (2019). Finally, MOSFiT indicates that the star was likely completely disrupted in the encounter, though the lower limit on the scaled β parameter b is consistent with a partial disruption when the systematic uncertainties are taken into account.

Table 4. MOSFiT TDE model parameter fits.

Quantity	Value	Units
$\log R_{\text{ph}0}$ (photosphere power-law constant)	$0.23^{+0.43}_{-0.42}$	–
$\log T_{\text{viscous}}$ (viscous delay time-scale)	$-0.09^{+0.39}_{-0.56}$	d
b (scaled impact parameter β)	$0.99^{+0.25}_{-0.93}$	–
$\log M_h$ (SMBH mass)	$6.89^{+0.22}_{-0.23}$	M_\odot
$\log \epsilon$ (efficiency)	$-0.44^{+0.71}_{-0.70}$	–
l (photosphere power-law exponent)	$1.84^{+0.27}_{-0.27}$	–
$\log n_{\text{H,host}}$ (local hydrogen column density)	$20.71^{+0.02}_{-0.02}$	cm^{-2}
M_* (stellar mass)	$0.10^{+0.37}_{-0.08}$	M_\odot
t_{exp} (time of disruption)	$-9.09^{+15.79}_{-15.86}$	d
$\log \sigma$ (model variance)	$-0.68^{+0.01}_{-0.01}$	–

Note. Best-fitting values and 1–99 per cent ranges for the MOSFiT TDE model parameters. Units are listed where appropriate. The listed uncertainties include both statistical uncertainties from the fit and the systematic uncertainties listed in table 3 of Mockler et al. (2019).

To understand peak emission in the context of stream-stream collisions, we use TDEMASS (Ryu, Krolik & Piran 2020b). TDEMASS assumes that the UV/optical emission is shock-powered and extracts the SMBH and stellar mass based on the observed peak luminosity and temperature at peak. Using our peak luminosity of $(6.14 \pm 0.17) \times 10^{44} \text{ erg s}^{-1}$ and temperature at peak of $50400^{+3800}_{-6300} \text{ K}$, we obtain a SMBH mass of $7.3^{+2.4}_{-1.0} \times 10^5 M_\odot$ and a disrupted stellar mass of $13.0^{+4.2}_{-2.0} M_\odot$, both consistent with the values obtained by (Ryu

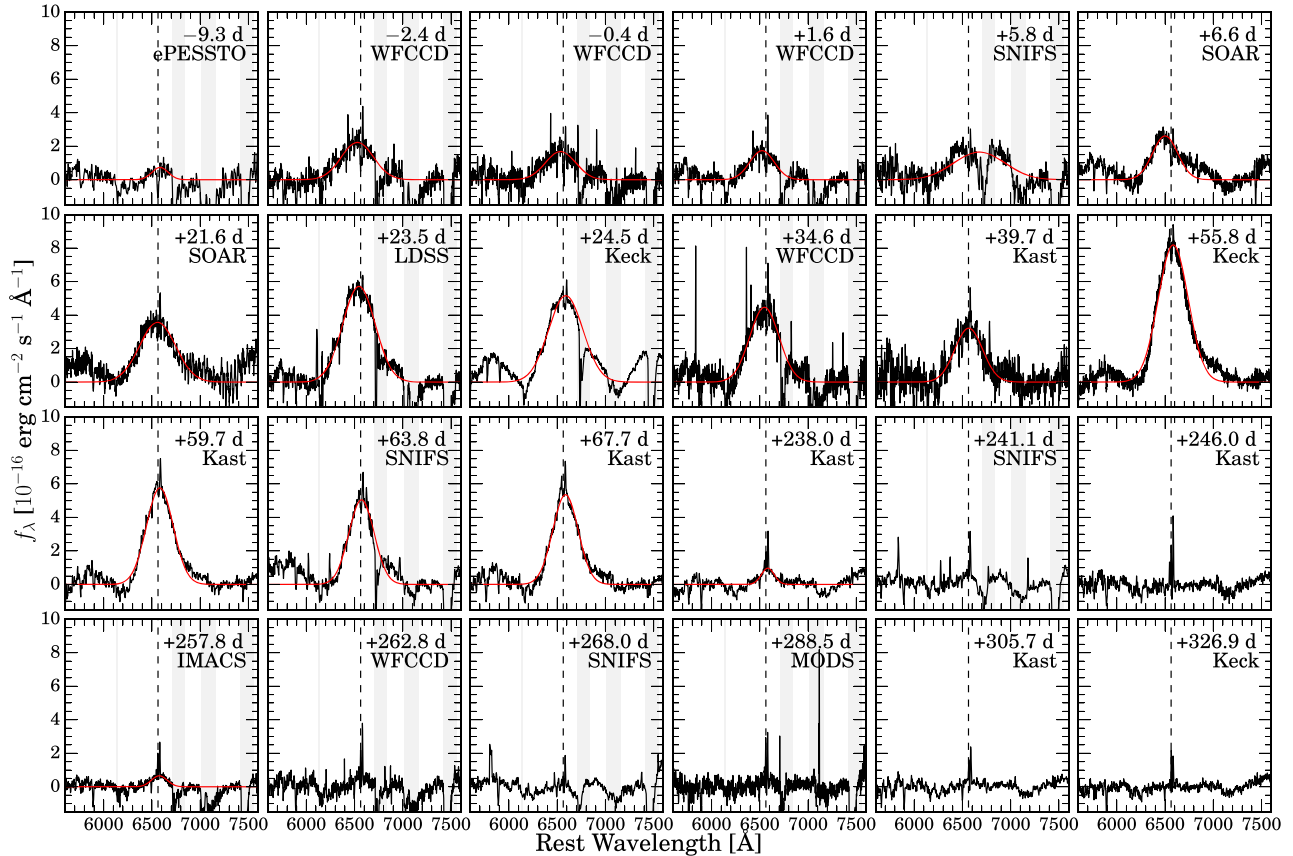


Figure 8. Spectral evolution of the $H\alpha$ line, increasing in time from top left to bottom right. The days relative to peak (MJD = 58548.5) in observer-frame days, of that particular spectral epoch and the instrument used to take the spectrum are shown in each individual panel. The red solid lines are Gaussian fits to the $H\alpha$ line profile, and are only shown for epochs with evidence of line emission. The vertical grey bands mark masked atmospheric telluric features.

et al. 2020b). Given the stellar mass of the host galaxy, this SMBH mass is roughly 15 times lower than expected from scaling relations. The stellar mass is unlikely for any TDE and inconsistent with the absence of recent star formation in the host galaxy.

3.2 Spectra

The early-time spectra of ASASSN-19dj have the very blue continuum that is a hallmark of tidal disruption events. This excess in blue flux grows towards peak light and fades back to host galaxy levels at later times. The very early-time optical spectra of ASASSN-19dj lack many of the spectral features that TDEs usually exhibit. For example, in the earliest spectrum, taken approximately 17 d after first light (see Section 3.1), there are no clear strong broad H and He lines.

We traced the evolution of the prominent $H\alpha$ feature by subtracting a linear continuum normalized in the regions around 6200 and 7200 Å. We then modelled the continuum-subtracted $H\alpha$ profile as a Gaussian. We masked narrow $H\alpha + [N II]$ lines for all spectra and telluric absorption features for spectra without telluric corrections. Looking at Fig. 8, which zooms into a small region around $H\alpha$, we see that the $H\alpha$ line slowly grows in strength from MJD ≈ 58539 to MJD ≈ 58544 , and quickly becomes very strong by MJD ≈ 58571 , with a peak line flux of $\sim 3.7 \times 10^{-13}$ erg cm $^{-2}$ s $^{-1}$ or a luminosity of $\sim 4.3 \times 10^{41}$ erg s $^{-1}$ at the distance of ASASSN-19dj. The line luminosity remains roughly constant until at least 41 d later, after which the source became Sun-constrained. After the seasonal gap,

we find evidence for weak $H\alpha$ emission as late as ~ 260 d after peak, consistent with other optical TDEs (e.g. Holoien et al. 2020; Hung et al. 2020).

Fig. 9 shows the luminosity and FWHM evolution of $H\alpha$ as a function of time and the luminosity versus line width. The line width is relatively constant prior to peak line flux and decreases thereafter, with some epochs exhibiting very broad $H\alpha$, up to roughly 2×10^4 km s $^{-1}$. There appears to be a positive correlation between $H\alpha$ line flux and line width. To test this, we performed the Kendall Tau test and find $\tau = 0.49$ and a corresponding p -value of 0.01 indicating a significant moderately strong correlation. The positive correlation seen here between $H\alpha$ line flux and line width agrees with what has been observed in the TDEs PS18kh (Holoien et al. 2019b; Hung et al. 2019), ASASSN-18pg (Leloudas et al. 2019; Holoien et al. 2020), and ASASSN-18ul (Wevers et al. 2019b). Unlike other TDEs such as ASASSN-18pg (Holoien et al. 2020), we do not see evidence for both a broad and narrow component of the $H\alpha$ emission.

While the dominant spectral features of ASASSN-19dj appear to be broad hydrogen lines, van Velzen et al. (2020) classify ASASSN-19dj as a TDE-Bowen object. There may be some evidence for broad emission centred on ~ 4600 Å, although the origin of this feature is difficult to determine. In order to analyse this feature in more detail, it was necessary to isolate the TDE emission flux. We first subtracted the host emission using the archival SDSS spectrum, then fit a Legendre polynomial to the continuum of the host-subtracted spectrum, and then subtracted the fitted continuum to get a host- and continuum-subtracted spectrum. The resulting spectra are shown in

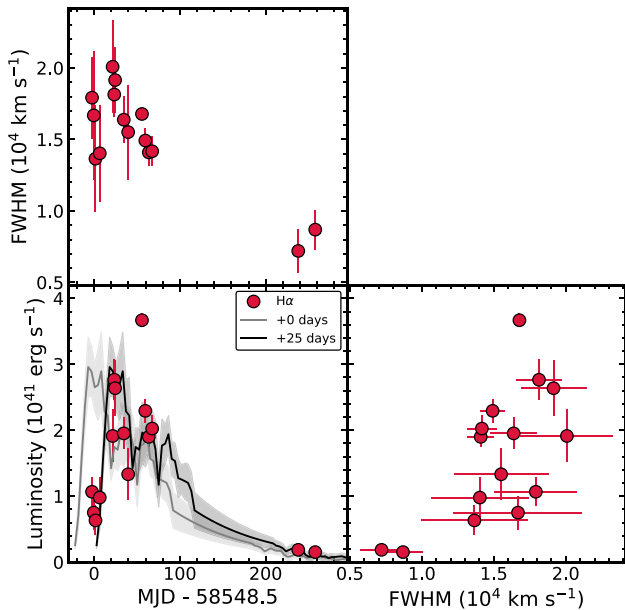


Figure 9. *Top left panel:* $H\alpha$ FWHM as a function of time. *Bottom left panel:* Time evolution of $H\alpha$ luminosity (red points). The bolometric light curve is scaled by a factor of 4.5×10^{-4} and shown with shaded error bars at zero days offset (grey) and 25 d offset (black) to highlight the delay in bolometric and $H\alpha$ peak. The bolometric light curves are smoothed by linearly interpolating to a time-series with the same length as the original coverage, but with half the number of points. *Right-hand panel:* $H\alpha$ luminosity as compared to the FWHM of the line.

Fig. 10 with the spectra selected to highlight the early temporal evolution.

By the +6.6 d SOAR spectrum, there is clear evidence for broad, blueshifted $H\alpha$, and $H\beta$, emission. The line profile of $H\alpha$ in particular is similar to PS17dzhz/AT2017eqx (Nicholl et al. 2019) and ZTF19abzrhgq/AT2019qiz (Nicholl et al. 2020), suggestive of an outflow (Roth & Kasen 2018). Interestingly, the $H\alpha$ and $H\beta$ lines in the +6.6 d +21.6 d SOAR spectra appear to have flat tops, like PS18kh (Holoien et al. 2019a) and ASASSN-18zj (Hung et al. 2020; Short et al. 2020), although for ASASSN-19dj this feature does not persist for long. Additionally, there is significant evolution in the broad features near 4600 Å. In pre-peak spectra, there is a very broad feature, similar to the early broad He II lines of ASASSN-15oi (Holoien et al. 2016a), which quickly change to distinct H lines. Similar to ZTF19abzrhgq/AT2019qiz, as the H lines return to their rest wavelengths the strengths of the Bowen fluorescence lines grow dramatically, with strong N III and He II emission by roughly 60 d after peak in addition to the now-dominant hydrogen emission.

Throughout the evolution of ASASSN-19dj, the $H\alpha$ line is very broad. There may be weak evidence for a broadening of the $H\alpha$ line at early times, similar to PS18kh (Holoien et al. 2019a), although these epochs have large uncertainties on the FWHM. The $H\alpha$ line begins to narrow again after the peak $H\alpha$ luminosity. There is a time delay between the peak UV/optical magnitude, which occurs at MJD = 58545.5, and the peak $H\alpha$ luminosity, at roughly MJD \simeq 58571. As can be seen in Fig. 9, the $H\alpha$ luminosity tracks the bolometric luminosity reasonably well when offset by 25 d. While this delay of \sim 25 d is only approximate, given our lack of spectra near this time, this provides an upper limit on how far from the SMBH this emission is located. If the lines are due to reprocessing of high-energy FUV and X-ray photons produced in an accretion disc, this suggests

the existence of reprocessing material at a distance of several tens of thousands of gravitational radii, slightly larger than the distance to reprocessing material derived for ASASSN-18pg, which showed clear evidence for a delay between the bolometric peak and peak line emission (Holoien et al. 2020). The early-time $H\alpha$ line profiles, with a blueshifted core and asymmetric red wing support the idea of an optically thick outflow as the reprocessing material (Roth & Kasen 2018) as do the existence of early-time He emission without corresponding H emission (Roth et al. 2016; Roth & Kasen 2018).

3.3 Spectral energy distribution

Fig. 11 shows the blackbody model fits in terms of luminosity, radius, and temperature for ASASSN-19dj. ASASSN-19dj is one of the most luminous TDEs discovered to date, with a peak luminosity of $(6.15 \pm 0.17) \times 10^{44}$ erg s $^{-1}$, consistent with the value derived by Liu et al. (2019). This peak luminosity is comparable only to the TDE ASASSN-18ul (Wevers et al. 2019b; Payne et al. in preparation) and the TDE/AGN ASASSN-18jd (Neustadt et al. 2019). The decline in bolometric luminosity of ASASSN-19dj is quite slow, which is consistent with the findings of Hinkle et al. (2020) that more luminous TDEs decay slower than less luminous TDEs. At later times (over \sim 250 d after peak) the luminosity appears to flatten out, consistent with other TDEs with late-time observations including ASASSN-14li (Brown et al. 2017), ASASSN-15oi (Holoien et al. 2016a), ASASSN-18pg (Leloudas et al. 2019; Holoien et al. 2020), ASASSN-18ul (Wevers et al. 2019b; Payne et al., in preparation), and ATLAS18way (van Velzen et al. 2020). In the cases of both ASASSN-19dj and ASASSN-15oi (Gezari et al. 2017; Holoien et al. 2018), the flattening of the bolometric light curve is roughly coincident with an increase in X-ray flux.

The blackbody radius of ASASSN-19dj is initially relatively small compared to other well-studied TDEs with similar strong $H\alpha$ emission such as ASASSN-18zj (Hung et al. 2020; Short et al. 2020), ASASSN-19bt (Holoien et al. 2019b), and PS18kh (Holoien et al. 2019a) as well as the TDE/AGN ASASSN-18jd (Neustadt et al. 2019). It is larger however than other X-ray bright TDEs like ASASSN-14li (Holoien et al. 2016b; Brown et al. 2017) and ASASSN-15oi (Holoien et al. 2016a). At late times the radius continues to decrease slowly, becoming consistent with other TDEs with well-sampled late-time evolution such as ASASSN-14li (Holoien et al. 2016b) and ASASSN-15oi (Holoien et al. 2018). Additionally this slow late-time decrease in radius is consistent with many TDEs in the literature (e.g. Hinkle et al. 2020; van Velzen et al. 2020).

Unlike other TDEs with strong $H\alpha$ emission (ASASSN-18zj, ASASSN-19bt, and PS18kh) and the TDE/AGN ASASSN-18jd, the temperature of ASASSN-19dj is quite hot, on the order of \sim 45 000 K. This temperature is more in line with the TDE-Bowen spectral class introduced by van Velzen et al. (2020). This hot temperature, especially at late times, is similar to the TDEs ASASSN-18pg, ASASSN-15oi, ASASSN-14li, with the latter two also exhibiting late-time X-ray emission. The blackbody temperatures of each of these TDEs are mostly flat throughout the evolution of the TDE. The blackbody temperature of ASASSN-19dj appears to decrease near peak, which can also be seen in Fig. 4 as the time of peak is earliest in the bluest bands and delayed in each of the red bands, similar to other TDEs with high-cadence pre-peak photometry such as ASASSN-19bt (Holoien et al. 2019b) and ASASSN-18pg (Holoien et al. 2020).

The UV/optical and X-ray SEDs at four epochs in the evolution of ASASSN-19dj are shown in Fig. 12. The UV/optical emission of ASASSN-19dj is dominant at early times and several orders of magnitude brighter than the X-ray emission. Using the SMBH mass

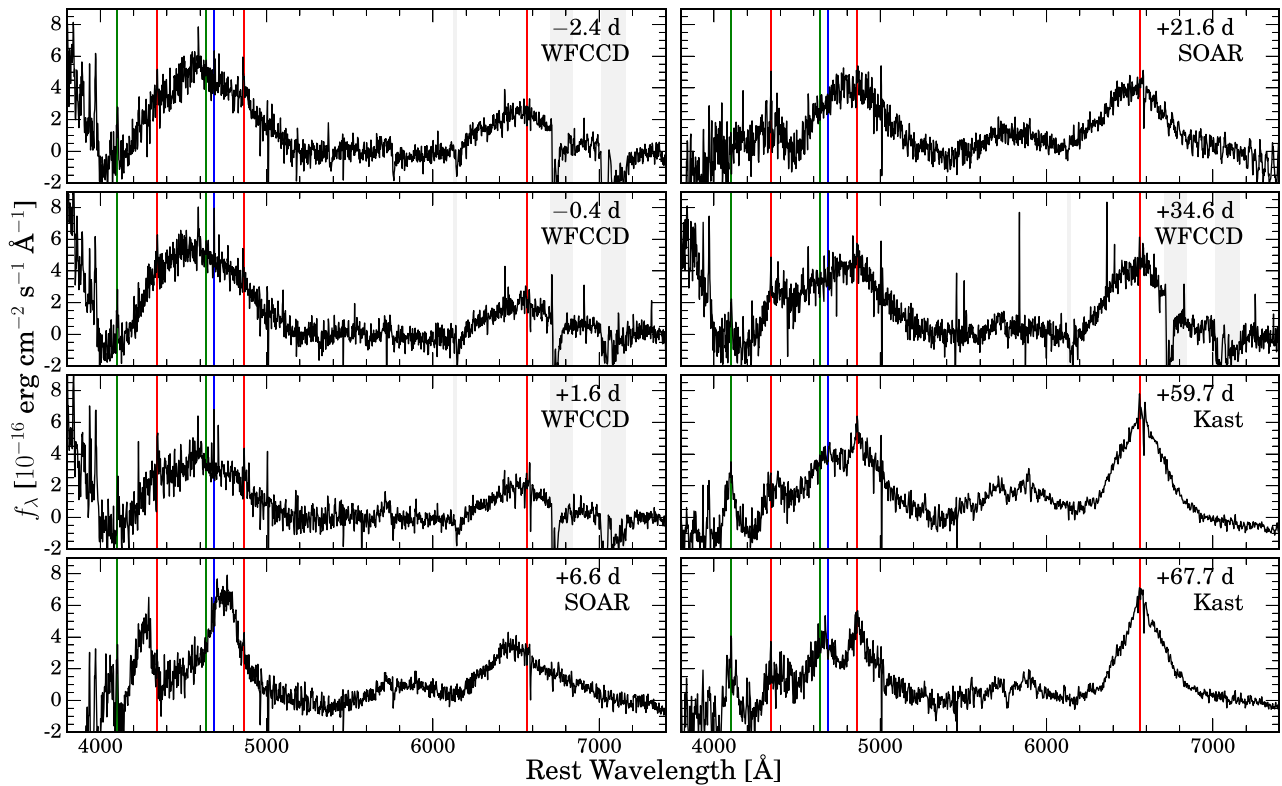


Figure 10. Host-subtracted spectral evolution of ASASSN-19dj increasing in time from top left to bottom right. The days relative to peak (MJD = 58548.5) in observer-frame days, of that particular spectral epoch and the instrument used to take the spectrum are shown in each individual panel. Vertical lines mark emission features common in TDEs with red marking H, blue marking He II, and green marking N III lines. The vertical grey bands mark masked atmospheric telluric features.

derived from MOSFiT, we calculate an Eddington luminosity of 9.8×10^{44} erg s $^{-1}$. Similar to what we find from the X-ray properties, we find no clear evidence for a plateau caused by Eddington-limited accretion near peak in either the bolometric light curve (see Fig. 11) or the single-band UV and optical light curves (see Fig. 4), although we caution that there is significant scatter in the SMBH mass estimates for this source. The Eddington ratio for the peak bolometric luminosity is ~ 0.6 , which is consistent with other UV/optical TDEs in the literature (e.g. Wevers et al. 2017; Mockler et al. 2019). By roughly 250 d after peak, the X-ray emission exceeded the UV/optical emission by roughly an order of magnitude, similar to the late-time X-ray brightening of ASASSN-15oi (Gezari et al. 2017; Holoiien et al. 2018), as well as the flatter late-time X-ray emission of ASASSN-14li (Brown et al. 2017). During the epoch of peak X-ray emission, at MJD ≈ 58782 , the difference between the UV/optical and X-ray SEDs is less pronounced. By roughly a year after peak, both the X-ray and UV/optical SEDs have faded and are comparable in peak luminosity.

3.4 Pre-ASASSN-19dj outburst?

The CRTS light curve of the host galaxy, KUG 0810+227, shows evidence of a previous outburst in September 2005 (see Fig. 3), roughly 14.5 yr prior to this TDE. The data quality of the archival CRTS images of the host galaxy were too poor to perform image subtraction. Instead, we stacked six CRTS images of the host galaxy during the outburst and seventeen reference images taken at least 5 yr after the outburst. Through comparison of the image centroids

for these stacks, we find the difference to be 0.18 ± 0.32 arcsec, corresponding to a physical distance of 84 ± 152 pc. Given the low quality of the archival images, this uncertainty is estimated by taking the standard deviation of the centroids of each of the individual images. While this constraint on the location of this previous transient is based on unsubtracted images, and therefore includes host light, it appears to be consistent with the host nucleus. However, we can make several statistical statements based on previous analysis presented in this paper.

The absolute magnitude of the brightest CRTS epoch is $V = -19.1$ or $L_V = 1.4 \times 10^{43}$ erg s $^{-1}$ (corrected for Galactic extinction, but assuming no host galaxy reddening), which is more luminous than the observed magnitudes of many types of supernovae, but consistent with the absolute magnitudes of Type Ia supernovae (e.g. Folatelli et al. 2010; Richardson et al. 2014). We attempted to fit the CRTS light curve of the archival outburst with SNOOPy (Burns et al. 2011) to constrain the properties of the light curve. We used the default $E(B - V)$ model, but assumed no host galaxy reddening given the single filter light curve. These fits get their shape information from the fitted Δm_{15} and use the K -corrections of Hsiao et al. (2007), the Milky Way dust map of Schlegel, Finkbeiner & Davis (1998), and SN templates of Prieto, Rest & Suntzeff (2006). From this, we find that the decline of this outburst is somewhat slower than expected for a Type Ia supernova, with $\Delta m_{15} = 0.8 \pm 0.4$ mag. Yet, given the data quality and maximum observed V -band magnitude, we cannot rule out a luminous Type Ia SN. The true peak luminosity of this transient is likely higher than 1.4×10^{43} erg s $^{-1}$, because a seasonal gap occurred immediately prior to the CRTS detection. Even if the

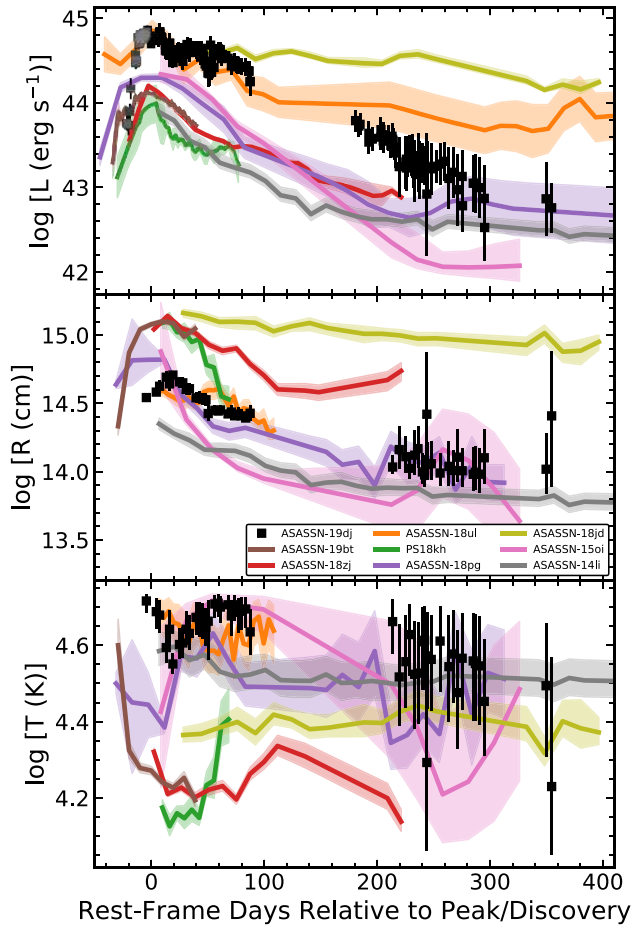


Figure 11. Evolution of the UV/optical blackbody luminosity (top panel), radius (middle panel), and temperature (bottom panel) for ASASSN-19dj (black squares), in comparison to other TDEs with strong $H\alpha$ emission, strong X-ray emission, or both: ASASSN-19bt (brown line; Holoien et al. 2019b), ASASSN-18zj (red line; van Velzen et al. 2020), ASASSN-18ul (orange line; Wevers et al. 2019b; Payne et al., in preparation), PS18kh (green line; Holoien et al. 2019a), ASASSN-18pg (purple line; Holoien et al. 2020), ASASSN-15oi (pink line; Holoien et al. 2018), and ASASSN-14li (grey line; Brown et al. 2017) in addition to the TDE/AGN ASASSN-18jd (gold line; Neustadt et al. 2019). The lines are smoothed over the individual epochs by linearly interpolating to a time-series with the same length as the original coverage, but with half the number of points. Time is in rest-frame days relative to the peak luminosity for the objects discovered prior to peak (ASASSN-19dj, ASASSN-19bt, ASASSN-18zj, ASASSN-18ul, PS18kh, and ASASSN-18pg), and relative to discovery for those which were not (ASASSN-18jd, ASASSN-15oi, and ASASSN-14li). The grey squares for ASASSN-19dj indicate where data has been bolometrically corrected using the ASAS-SN g-band light curve assuming the temperature prior to the first *Swift* epoch was constant.

peak luminosity is higher than this it may still be consistent with the tail of observed SN Ia magnitudes (e.g. Folatelli et al. 2010; Richardson et al. 2014) or a superluminous supernova.

Next, we evaluated the possibility that this CRTS flare was a previous TDE, by estimating a TDE rate for the host galaxy. The rate of TDEs is roughly 10^{-4} to 10^{-5} yr $^{-1}$ per galaxy (e.g. van Velzen & Farrar 2014; Holoien et al. 2016b; Auchettl et al. 2018) for an average galaxy. However, KUG 0810+227 is a post-starburst galaxy, for which it is known that the TDE rate can be enhanced by up to 200 times the average (e.g. French et al. 2016; Law-Smith et al.

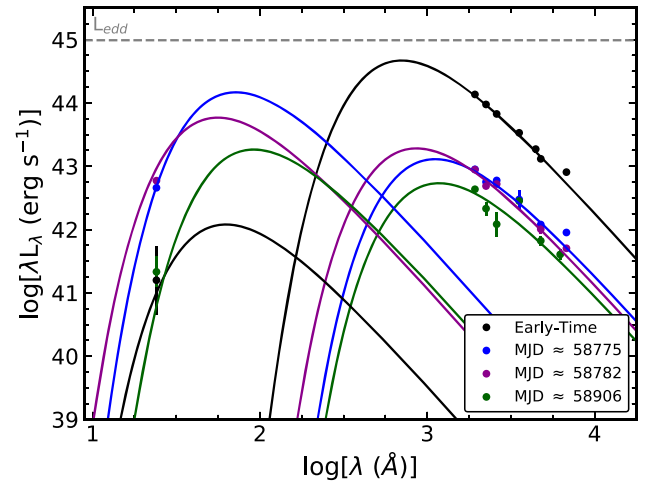


Figure 12. Spectral energy distribution of ASASSN-19dj at four different epochs. The first epoch (black) is at MJD = 58544 for the UV/optical emission and is the stacked early-time *Swift* for the X-ray. The second epoch (blue) is the first epoch after the X-ray brightening at MJD \approx 58775 for both the UV/optical and X-ray. The third epoch (purple) is the epoch of peak X-ray emission at MJD \approx 58782. The fourth epoch (green) is a late-time epoch at MJD \approx 58906. For each epoch, the data is shown as points while the lines represent the best-fitting blackbody components for the UV/optical and X-ray emission. The dashed grey line indicates the Eddington luminosity for an SMBH of mass $7.8 \times 10^6 M_\odot$, the SMBH mass derived from MOSFIT.

2017; Graur et al. 2018), and thus it would not be unreasonable that a TDE could occur every 50–500 yr. Even within this sample of post-starburst galaxies, KUG 0810+227 appears to be extreme in terms of its Lick $H\delta_A$ index. Counting galaxies with $H\delta_A - \sigma(H\delta_A) > 7.0 \text{ \AA}$ and $H\alpha$ emission EW $< 3.0 \text{ \AA}$, gives just 0.025 per cent of all the galaxies in SDSS and a similar TDE rate enhancement of ~ 250 times the average. Thus, there is the possibility that the CRTS flare is a previous TDE.

Several pieces of archival data and optical emission line diagnostics are consistent with KUG 0810+227 being a LLAGN. The line ratios of the optical spectrum of KUG 0810+227 lie in the Seyfert region of two line ratio diagnostic diagrams, which suggests the possibility of the host being an LLAGN. However, we note that the WHAN diagram classifies KUG 0810+227 as an RG, suggesting a possible non-AGN ionization source. In the X-ray, the first *Swift* XRT epoch gives a deep upper limit, which is consistent with a small fraction of observed X-ray luminosities of AGN (e.g. Tozzi et al. 2006; Ricci et al. 2017). Therefore, because the host is consistent with an LLAGN, the previous flare could be associated with a pre-flare AGN outburst.

3.5 X-rays

ASASSN-19dj is one of several optical TDEs to show strong X-ray emission. In Fig. 13, we show the stacked early-time *Swift* spectrum and an *NICER* spectrum from the epoch of peak X-ray emission. In Fig. 14 (top panel), we show the X-ray light curve as derived from both the individual *Swift* and *NICER* observations. To estimate the X-ray luminosity, we converted the extracted count rate into flux using WEBPIMMS¹¹ and assumed an absorbed blackbody emission model with a temperature of ~ 50 eV, corresponding to the average

¹¹<https://heasarc.gsfc.nasa.gov/cgi-bin/Tools/w3pimms/w3pimms.pl>

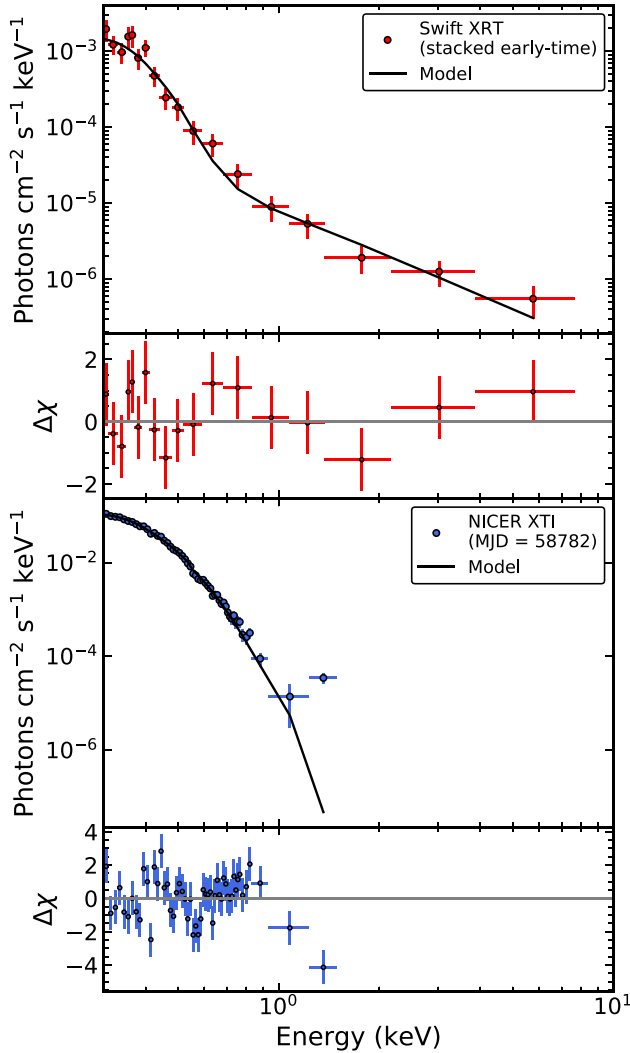


Figure 13. X-ray spectrum and $\Delta\chi$ for the stacked early-time *Swift* observations (top panel) and *NICER* observations of ASASSN-19dj at peak X-ray emission on MJD = 58782 (bottom panel). The black lines in the top and bottom panels are the best-fitting absorbed blackbody model + power-law and absorbed blackbody models, respectively. The reduced χ^2 of the best-fitting models are 0.90 for 15 degrees of freedom and 1.72 for 50 degrees of freedom, respectively.

blackbody temperature derived from our *Swift* and *NICER* X-ray spectra. This value is also consistent with the blackbody temperatures of other X-ray bright TDEs (e.g. ASASSN-14li, ASASSN-15oi, Brown et al. 2017; Holoien et al. 2018; Kara et al. 2018). The first X-ray observation of ASASSN-19dj was taken using the *Swift* XRT approximately 4 d before the peak UV/optical emission (MJD = 58544.8). During this observation, ASASSN-19dj showed no evidence of X-ray emission with a 3σ upper limit of 6×10^{40} erg s $^{-1}$, consistent with the limits/detection of X-ray emission seen prior to peak in the TDEs ASASSN-19bt (Holoien et al. 2019b), ASASSN-18pg (Leloudas et al. 2019; Holoien et al. 2020), ZTF19abzrhgq/AT2019qiz (Auchettl et al. 2019) and other X-ray TDE candidates (Auchettl et al. 2017). This upper limit places even stricter constraints on the possibility that the host galaxy is an LLAGN, with Tozzi et al. (2006) finding that fewer than ~ 10 per cent of AGN have X-ray luminosities this low, and Ricci et al. (2017)

measuring only 1 per cent of their unobscured non-blazar AGN sample to have X-ray luminosities this low.

ASASSN-19dj was first detected in X-rays ~ 9 d later, ~ 4 d after the UV/optical peak, in the second *Swift* XRT observation, with its X-ray luminosity increasing by at least half an order of magnitude to $\sim 3 \times 10^{41}$ erg s $^{-1}$. Similar to ASASSN-18jd (Neustadt et al. 2019), ASASSN-18ul (Wevers et al. 2019b), and ASASSN-15oi (Gezari et al. 2017), the X-ray emission of ASASSN-19dj showed significant variations in luminosity over the first ~ 100 d after peak, varying between $\sim 10^{40.7}$ and $10^{41.7}$ erg s $^{-1}$ before the seasonal gap, much larger than the variability seen in ASASSN-14li (Brown et al. 2017), but similar to that seen in ASASSN-18jd (Neustadt et al. 2019) or ASASSN-18ul (Wevers et al. 2019b). Once the source became visible again ~ 220 d after peak, *XMM-Newton* slew, *Swift* XRT, and *NICER* observations found that the source had brightened by nearly a factor of ~ 10 . This brightening behaviour is reminiscent of what was seen in ASASSN-15oi (Gezari et al. 2017; Holoien et al. 2018) and hinted at in ASASSN-18ul (Wevers et al. 2019b), in which the X-ray emission increased by an order of magnitude ~ 250 d after peak brightness before fading. ASASSN-19dj peaked at an X-ray luminosity of $\sim 10^{43}$ erg s $^{-1}$ before fading by nearly an order of magnitude over ~ 100 d and then plateauing at an X-ray luminosity of $\sim 10^{42}$ erg s $^{-1}$. The peak luminosity corresponds to an Eddington ratio between 0.01 and 0.03, consistent with other X-ray bright TDEs (Mockler et al. 2019; Wevers et al. 2019a) and again disfavouring Eddington-limited accretion as suggested by van Velzen et al. (2019).

In Fig. 14 (second panel), we present the evolution of the X-ray hardness ratio¹² (HR) as a function of time. At early times, ASASSN-19dj shows significant variability in its hardness, varying between a soft HR of -1 and harder HR of -0.2 during the first 100 d. By 200–280 days after days after peak, the hardness ratio of ASASSN-19dj plateaued to a soft HR between -1 and -0.8 , before hardening significantly over ~ 20 d from 280 to 300 d after peak. Finally, from 300 d after peak onwards, ASASSN-19dj returned to the HR variability observed at early times. The behaviour seen at early times is consistent with the presence of hard X-ray emission in the form of a power law in addition to a soft thermal blackbody, consistent with Liu et al. (2019). This can be seen in the merged *Swift* spectra derived from the observations taken within the first 100 d (see Fig. 13). The softening of the X-ray emission between 200 and ~ 280 d after discovery occurs when the X-ray emission from this event becomes dominated by a strong thermal blackbody component. Near peak, the constant HR with time and decreasing X-ray luminosity, is consistent with that exhibited by non-thermal TDEs such as ASASSN-14li (Auchettl et al. 2018). However, we note that the lack of significant HR evolution seen in ASASSN-14li begins at peak brightness in X-ray, UV/optical, and bolometric luminosity and continues for thousands of days after peak. As the blackbody component cools with time and fades, the X-ray emission is seen to harden, similar to what was seen in ASASSN-14li (e.g. Kara et al. 2018).

During the evolution of ASASSN-19dj, the HR and the X-ray luminosity seem to follow an inverse relationship, where the X-ray emission becomes harder as the luminosity of the source fades, and becomes softer as the source brightens. This relationship is shown in Fig. 15, with colour-coding and arrows to highlight the trend. This evolution is consistent with what is seen in highly variable, X-ray bright AGNs (cf. Fig. 4 of Auchettl et al. 2018). Additionally this

¹²The hardness ratio (HR) is defined as $HR = (H - S)/(H + S)$, where H is the number of counts in the 2.0–10.0 keV energy range and S is the number of counts in the 0.3–2.0 keV energy range.

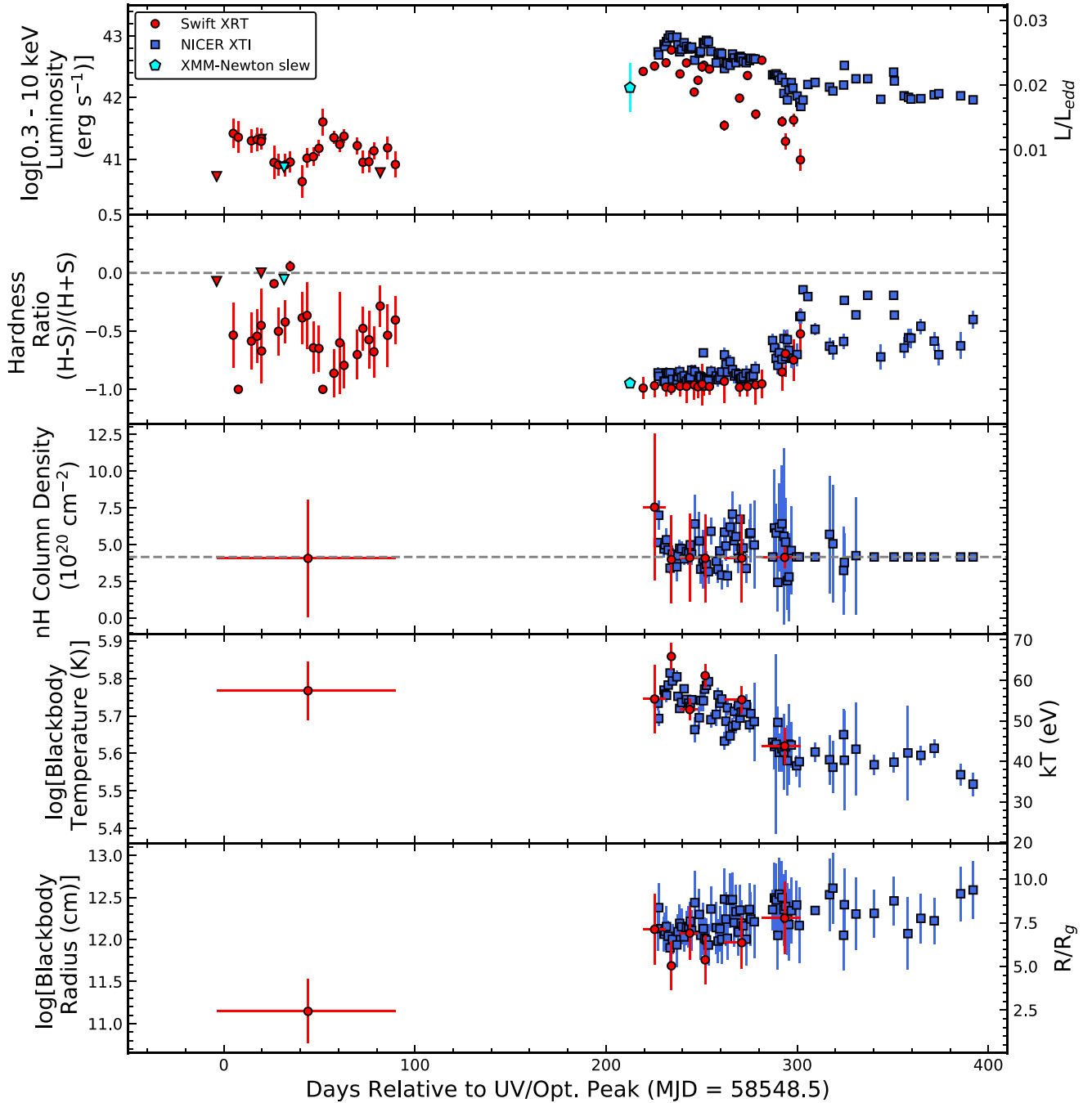


Figure 14. In order from top to bottom: X-ray luminosity, hardness ratio, neutral hydrogen column density (with the dashed grey line marking the Galactic column density), blackbody temperature, and blackbody radius of ASASSN-19dj measured with *Swift* (red circles), *NICER* (blue squares), and *XMM-Newton* slew (cyan pentagon). We define hard counts H as the number of counts in the 2–10 keV range and soft counts S are the number of counts in the 0.3–2 keV range, with a grey dashed line marking zero. The hardness ratio is defined as $(H - S)/(H + S)$. Downward-facing triangles mark upper limits.

behaviour is similar to the overall trends between X-ray luminosity and spectral hardness seen in the sample of Wevers (2020). The overall behaviour seen in ASASSN-19dj is quite unique compared to all other X-ray TDE candidates, even compared to ASASSN-15oi, which showed delayed brightening of the X-ray emission ~ 200 d after peak UV brightness (Gezari et al. 2017), or ASASSN-18jd which showed large variations in HR with time before the emission completely faded (Neustadt et al. 2019). The correlated changes in HR and luminosity have not been seen before and the late-time

brightening for ASASSN-19dj is different than that of ASASSN-15oi. Auchettl et al. (2018) showed that <4 per cent of X-ray bright AGN could produce flare emission that exhibits a coherent decay and a constant HR similar to that of an X-ray bright TDE. So while the brightening is similar to what we see from thermal TDEs such as ASASSN-14li, we cannot rule out that some of the emission arises from a pre-existing AGN disc (e.g. Blanchard et al. 2017).

To further explore the nature of the X-ray emission arising from ASASSN-19dj, we analysed the *Swift* and *NICER* spectra using the

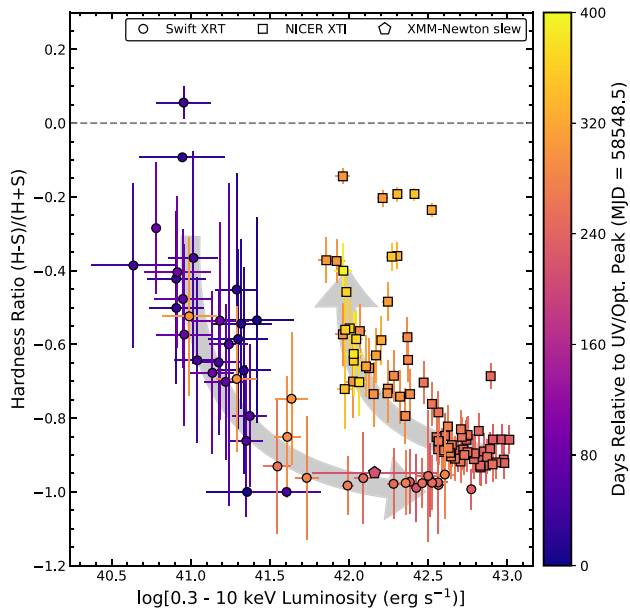


Figure 15. Hardness ratio as a function of X-ray luminosity as measured by *Swift* (circles), *NICER* (squares), and *XMM-Newton* slew (pentagon) for the detections only. Light grey arrows indicate that as ASASSN-19dj becomes brighter, the X-ray emission becomes softer. This behaviour is similar to that seen in X-ray bright AGNs (Auchettl et al. 2018). The colour bar on the right indicates the phase relative to the UV/optical peak, with darker colours indicating earlier times.

X-ray spectral fitting program XSPEC version 12.10.1f (Arnaud 1996), and chi-squared statistics. While we fit the majority of *NICER* spectra individually (with the exception of a handful of observations at late times), it was necessary to stack the early-time *Swift* observations to get adequate S/N. We show the results of these spectral fits in the bottom three panels of Fig. 14.

At early times, the merged *Swift* spectrum is best fit by an absorbed blackbody plus power-law model. However, at late times, when ASASSN-19dj is significantly brighter, we find that an absorbed blackbody is sufficient to model the observed spectra. We let the column density (N_H), blackbody temperature (kT) and blackbody normalization, as well as the photon index Γ and power-law normalization for the early *Swift* spectra, of each model be free parameters. In Table 5, we summarize the best-fitting parameters of our spectral fits.

Table 5. X-ray spectral parameters of ASASSN-19dj.

MJD	N_H (10^{20} cm^{-2})	N_H error (10^{20} cm^{-2})	kT (keV)	kT error (keV)	log radius (cm)	Radius error (cm)	Satellite
58544.75–58638.27	4.06	4.00	0.050	0.009	11.15	0.38	<i>Swift</i>
58767.93–58779.88	7.55	5.00	0.048	0.010	12.12	0.42	<i>Swift</i>
58782.67	3.98	3.00	0.062	0.005	11.69	0.29	<i>Swift</i>
...
58920.19	4.16	–	0.035	0.002	12.22	0.27	<i>NICER</i>
58934.02	4.16	–	0.030	0.002	12.54	0.33	<i>NICER</i>
58940.53	4.16	–	0.028	0.002	12.59	0.34	<i>NICER</i>

Note. Neutral hydrogen column densities, blackbody temperatures, and effect blackbody radii derived from the various X-ray spectral epoch. A range of MJD in the first column indicates the beginning and end of the range over which data were stacked to increase S/N. The last column reports the source of the data for each epoch. Only a small section of the table is displayed here. The full table can be found online as an ancillary file.

The third panel of Fig. 14 shows the column density as a function of time. The column densities derived using the *NICER* spectra and the merged *Swift* spectra are all consistent with the Galactic column density along the line of sight, although the uncertainties are large.

In Fig. 14 (fourth panel), we show the temperature evolution of ASASSN-19dj. We find that the derived X-ray blackbody temperatures are similar to other X-ray bright TDEs such as ASASSN-14li (Holoien et al. 2016b; Brown et al. 2017) and ASASSN-15oi (Gezari et al. 2017; Holoien et al. 2018), and the very tail end of the blackbody temperature distribution of unobscured non-blazar AGNs (Ricci et al. 2017), peaking at kT \sim 110 eV. We find that the temperature of ASASSN-19dj is lower than derived for the TDE/AGN candidates ASASSN-18jd (Neustadt et al. 2019) and ASASSN-18ul (Wevers et al. 2019b), which had blackbody temperatures more consistent with known AGN. Initially, we find that the blackbody component had a mean temperature of \sim 60 eV. Unfortunately, due to the faintness of the source at early times, we are unable to constrain whether the temperature is constant with time or varies as seen during the late phases of its evolution. Interestingly, when the luminosity of the source increases after the seasonal gap, we find that the temperature initially does not change significantly from that seen at early times. However, as the source increases to peak brightness, we find that the temperature also increases, peaking at 62 ± 5 eV. As the source begins to fade, the temperature seems to follow the same short time-scale variability behaviour seen in the X-ray light curve, suggesting that the short time-scale luminosity variation we observe is dominated by changes in the blackbody temperature with time. The change in temperature is most dramatic between 200 and 280 d after the peak of the UV/optical light curve, where the blackbody temperature drops from \sim 60 to \sim 40 keV, before plateauing at this lower temperature value for the next \sim 100 d, similar to what was seen in ASASSN-14li after peak X-ray brightness (see table 3 of Brown et al. 2017).

In the bottom panel of Fig. 14, we show the evolution of the effective blackbody radius as a function of time. During its early evolution ASASSN-19dj has a blackbody radius that is consistent with ASASSN-18jd (Neustadt et al. 2019) and ASASSN-15oi (Holoien et al. 2018) both before and after its observed X-ray brightening. The brightening of ASASSN-19dj after the seasonal gap is also associated with a dramatic order of magnitude increase in the blackbody radius. This suggests that the increase in X-ray luminosity is a result of an expansion of the X-ray emitting region rather than delayed accretion that could result from inefficient circularization as suggested for ASASSN-15oi (Gezari et al. 2017) and previously suggested for ASASSN-19dj (Liu et al. 2019). van Velzen et al. (2020) also find that the X-ray brightening of ASASSN-19dj does not require delayed

Table 6. Spectroscopic observations of ASASSN-19dj.

MJD	Date	Telescope	Instrument	Rest wavelength range (Å)	Exposure time (s)
58539.2	2019 February 25.2	ESO New Technology Telescope 3.58-m	EFOSC2	3555–9027	1 × 300
58546.1	2019 March 4.1	du Pont 100-in	WFCCD	3717–9390	1 × 900
58548.1	2019 March 6.1	du Pont 100-in	WFCCD	3717–9390	3 × 900
58550.1	2019 March 8.1	du Pont 100-in	WFCCD	3717–9390	1 × 900
58554.3	2019 March 12.3	University of Hawaii 88-in	SNIFS	3229–9489	3 × 1800
58555.1	2019 March 13.1	Southern Astrophysical Research Telescope 4.1-m	GHTS	3913–8701	1 × 600
58570.1	2019 March 28.1	Southern Astrophysical Research Telescope 4.1-m	GHTS	3913–8701	1 × 450
58572.0	2019 March 30.0	Magellan Clay 6.5-m	LDSS-3	3619–9048	1 × 400
58573.3	2019 March 31.3	Keck I 10-m	LRISp	3130–9781	4 × 1800
58583.1	2019 April 10.1	du Pont 100-in	WFCCD	3717–9390	3 × 900
58588.2	2019 April 15.3	Lick Shane Telescope 120-in	Kast	3067–10039	1 × 1545 (blue) 3 × 500 (red)
58604.3	2019 May 1.3	Keck I 10-m	LRIS	4500–9709	1 × 190 (blue) 1 × 180 (red)
58608.2	2019 May 5.2	Lick Shane Telescope 120-in	Kast	3067–10039	1 × 1545 (blue) 3 × 500 (red)
58612.3	2019 May 9.3	University of Hawaii 88-in	SNIFS	3229–9489	2 × 1800
58616.2	2019 May 13.2	Lick Shane Telescope 120-in	Kast	3067–10039	1 × 1230 (blue) 2 × 600 (red)
58786.5	2019 October 30.5	Lick Shane Telescope 120-in	Kast	3067–10039	2 × 1230 (blue) 4 × 600 (red)
58789.6	2019 November 2.6	University of Hawaii 88-in	SNIFS	3229–9489	2 × 2100
58794.5	2019 November 7.5	Lick Shane Telescope 120-in	Kast	3067–10039	1 × 1865 (blue) 3 × 600 (red)
58806.3	2019 November 19.3	Magellan Baade 6.5-m	IMACS	4157–9195	4 × 900
58811.3	2019 November 24.3	du Pont 100-in	WFCCD	3717–9390	1 × 1200
58816.5	2019 November 29.5	University of Hawaii 88-in	SNIFS	3229–9489	1 × 3600
58839.3	2019 December 22.3	Large Binocular Telescope 8.4 m	MODS	3130–9781	4 × 1200
58854.2	2020 January 6.2	Lick Shane Telescope 120-in	Kast	3067–10039	2 × 1060 (blue) 3 × 700 (red)
58875.4	2020 January 27.4	Keck I 10-m	LRIS	3063–9709	1 × 900 (blue) 1 × 888 (red)

Note. Modified Julian Day, calendar date, telescope, instrument, wavelength range, and exposure time for each of the spectroscopic observations obtained of ASASSN-19dj for the initial classification and during our follow-up campaign.

accretion. The size of the blackbody radius of ASASSN-19dj is consistent with ASASSN-14li (Brown et al. 2017), ASASSN-15oi (Holoien et al. 2018), and other TDE candidates whose blackbody radii were measured at peak using X-rays (see fig. 11 of Wevers et al. 2019a).

Similar to Wevers et al. (2019a) and individual studies of X-ray bright TDE candidates (e.g. Brown et al. 2017; Holoien et al. 2018), we find that the X-ray emitting region is at least an order of magnitude smaller than the blackbody radius interfered from the UV/optical (see Fig. 11). Additionally, the blackbody radius is smaller than (at early times) and equal to (at peak brightness) the ISCO (innermost stable circular orbit) of the black hole, assuming a Schwarzschild black hole. Due to the high cadence and high collecting area of *NICER*, we are also able to observe how the blackbody radius changes as a function of time. We find that the blackbody radius does not vary significantly over short time-scales and shows a slow overall increase with time. The slow evolution at late times is consistent with that seen in ASASSN-14li (Brown et al. 2017) and ASASSN-15oi (Holoien et al. 2018).

4 DISCUSSION

ASASSN-19dj has one of the most complete data sets of any TDE to date. Its temporal coverage stretches from early in its evolution to more than a year after peak and spans the X-ray, NUV, and optical energy bands. As such, ASASSN-19dj provides us with the opportunity to make comparisons to the various theoretical models put forth to explain the UV/optical emission of TDEs. As described previously, the two main models are reprocessed disc emission (e.g. Dai et al. 2018; Mockler et al. 2019) and stream–stream collisions (e.g. Jiang et al. 2016; Bonnerot et al. 2017; Lu & Bonnerot 2020; Ryu et al. 2020a). Since one of the major contentions of the reprocessing picture of TDE emission is that material surrounding the SMBH absorbs X-ray photons and re-emits them at longer wavelengths, our

combination of a rich set X-ray and UV/optical observations is well suited to test these predictions.

A key feature of the evolution of ASASSN-19dj that may help distinguish between the models is the late-time X-ray brightening. Liu et al. (2019) suggest this is evidence for delayed accretion, similar to ASASSN-15oi (Gezari et al. 2017). ASASSN-19dj, however shows significant X-ray emission within days of the UV/optical peak. Given the deep upper limit on the X-ray emission from the first *Swift* epoch ($<6 \times 10^{40}$ erg s $^{-1}$), the early X-ray emission must be associated with the ASASSN-19dj TDE rather than any pre-existing AGN activity. This is evidence for prompt circularization and the formation of an accretion disc. Additionally, the hardness–luminosity correlations shown in Fig. 15 are consistent with variable accretion on to a compact object, further supporting the early-time creation of a disc.

Another possible explanation for the X-ray brightening is a changing obscuration, which allows us to see more of the dominant soft X-ray emission. There are two issues with this however. First, the stacked early-time *Swift* X-ray spectrum shows clear evidence of strong soft X-ray emission. Secondly, the column density inferred from the early-time *Swift* spectrum shows no evidence of being elevated relative to the column density at late times. The column densities are modest ($\sim 5 \times 10^{20}$ cm $^{-2}$) and consistent with Galactic.

Our dense spectral coverage of ASASSN-19dj (detailed in Table 6) also allow us to place this source in the broader context of TDE flares. The broad feature near 4600 Å is commonly associated with Bowen fluorescence (Leloudas et al. 2019; Neustadt et al. 2019; van Velzen et al. 2020) and requires FUV flux for creation. In the spectral sequence of ASASSN-19dj, this broad complex of lines is present in the first pre-peak spectrum and persists until the seasonal break. However, as seen in Fig. 10, the makeup of this feature is complex and evolves rapidly from a broad hump to individual broad lines as the TDE fades. If an accretion disc is the dominant producer of FUV radiation, this would suggest that an accretion disc has been formed

very early on in the evolution of ASASSN-19dj. However, many of the emission lines do not peak until after the bolometric peak, with H α peaking roughly 2 months after peak UV/optical light. This delay may be explained if the emission is driven by stream–stream collisions (e.g. Piran et al. 2015; Krolik et al. 2016; Ryu et al. 2020a). We also note that the line profiles are consistent with theoretical predictions for outflows, including a delay in the appearance of strong Balmer emission (Roth & Kasen 2018).

5 SUMMARY

We have presented multiwavelength photometric and spectroscopic data of the tidal disruption event ASASSN-19dj. For the third time, we observe the initial optical rise of a TDE and find that it is again consistent with flux $\propto t^2$. ASASSN-19dj is among the most UV luminous TDEs yet discovered, with a peak absolute $UVW2$ magnitude of -21.01 ± 0.04 mag, as well as being among the hottest. The evolution of the UV/optical emission of ASASSN-19dj is roughly consistent with that of other TDEs, following the trend that more luminous TDEs decay more slowly after peak. Our set of 24 spectra follow the evolution of this TDE from 9 d before peak until 327 d after, showing significant changes in both the continuum and H α emission in this time period. The peak H α emission is delayed by roughly 25 d from the UV/optical luminosity peak.

Through a search of archival CRTS photometry, we find a previous flare in the host galaxy roughly 14.5 yr prior to ASASSN-19dj. While the quality of the CRTS images was poor, we were able to perform image centroiding and find that the location of the previous outburst is consistent with the nucleus. Given the fact that KUG 0810+227 is a post-starburst galaxy, we would expect the TDE rate to be significantly higher in this galaxy than an average galaxy, allowing for the possibility of an earlier TDE. However, given the available data, the earlier flare could also be a luminous nuclear supernova or some other form of accretion flare from the SMBH.

In addition to being luminous in the UV and optical, ASASSN-19dj increased in X-ray luminosity near peak UV/optical light. After the 2019 seasonal gap, the X-ray luminosity was observed by *XMM-Newton*, *Swift*, and *NICER* to have increased by an additional order of magnitude. The increase in X-ray luminosity appears to be a consequence of an increase in the area of the X-ray emitting region, while the short-term variability and late-time decrease seen in the X-ray light curve arise from changes in the X-ray temperature.

ASASSN-19dj is one of the few tidal disruption events with extensive multiwavelength photometric and spectroscopic data spanning from before peak to more than a year after. Even so, details on the emission mechanisms and spectral evolution are difficult to constrain. This indicates the importance of surveys like ASAS-SN and their ability to quickly find and confirm TDEs early so that similarly comprehensive data sets can be constructed.

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DATA AVAILABILITY

The data underlying this article are available in the article and in its online supplementary material.

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SUPPORTING INFORMATION

Supplementary data are available at [MNRAS](https://www.mnras.org) online.

Table 2. Host-subtracted photometry of ASASSN-19dj.

Table 3. X-ray luminosity and hardness ratios of ASASSN-19dj.

Table 5. X-ray spectral parameters of ASASSN-19dj.

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¹*Institute for Astronomy, University of Hawai'i, 2680 Woodlawn Dr., Honolulu, HI 96822, USA*

²*The Observatories of the Carnegie Institution for Science, 813 Santa Barbara St, Pasadena, CA 91101, USA*

³*DARK, Niels Bohr Institute, University of Copenhagen, Lyngbyvej 2, DK-2100 Copenhagen, Denmark*

⁴*Department of Astronomy and Astrophysics, University of California, Santa Cruz, CA 95064, USA*

⁵*School of Physics, The University of Melbourne, Parkville, VIC 3010, Australia*

⁶*ARC Centre of Excellence for All Sky Astrophysics in 3 Dimensions (ASTRO 3D), Sydney 2006, Australia*

⁷*Department of Astronomy, The Ohio State University, 140 West 18th Avenue, Columbus, OH 43210, USA*

⁸*Center for Cosmology and Astroparticle Physics, The Ohio State University, 191 W. Woodruff Avenue, Columbus, OH 43210, USA*

⁹*California Institute of Technology, 1200 E. California Blvd, Pasadena, CA 91125, USA*

¹⁰*European Southern Observatory, Alonso de Córdova 3107, Vitacura, Casilla 19001, Santiago, Chile*

¹¹*Kavli Institute for Astronomy and Astrophysics, Peking University, Yi He Yuan Road 5, Hai Dian District, Beijing 100871, China*

¹²*Millennium Institute of Astrophysics, 8320000 Santiago, Chile*

¹³*Departamento de Ciencias Físicas, Universidad Andres Bello, Avda. Republica 252, 8320000 Santiago, Chile*

¹⁴*Rubin Observatory Project Office, 950 North Cherry Avenue, Tucson, AZ 85719, USA*

¹⁵*Homer L. Dodge Department of Physics & Astronomy, The University of Oklahoma, 440 W. Brooks Street, Norman, OK 73019, USA*

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