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### Authors

Amanullah, R  
Johansson, J  
Goobar, A  
[et al.](#)

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# Diversity in extinction laws of Type Ia supernovae measured between 0.2 and 2 $\mu\text{m}$

R. Amanullah,<sup>1\*</sup> J. Johansson,<sup>1</sup> A. Goobar,<sup>1</sup> R. Ferretti,<sup>1</sup> S. Papadogiannakis,<sup>1</sup> T. Petrushevska,<sup>1</sup>, P. J. Brown,<sup>2</sup> Y. Cao,<sup>3</sup> C. Contreras,<sup>4</sup> H. Dahle,<sup>5</sup> N. Elias-Rosa,<sup>6</sup> J. P. U. Fynbo,<sup>7</sup> J. Gorosabel,<sup>8,9</sup> L. Guaita,<sup>10</sup> L. Hangard,<sup>1</sup> D. A. Howell,<sup>11,12</sup>, E. Y. Hsiao,<sup>13,4</sup> E. Kankare,<sup>14</sup> M. Kasliwal,<sup>15</sup> G. Leloudas,<sup>16,7</sup> P. Lundqvist,<sup>17</sup> S. Mattila,<sup>18</sup> P. Nugent,<sup>19,20</sup> M. M. Phillips,<sup>3</sup> A. Sandberg,<sup>17</sup> V. Stanishev,<sup>21</sup> M. Sullivan,<sup>22</sup> F. Taddia,<sup>17</sup> G. Östlin<sup>17</sup> S. Asadi,<sup>17†</sup> R. Herrero-Illana,<sup>8†</sup> J. J. Jensen,<sup>7†</sup> K. Karhunen,<sup>23†</sup> S. Lazarevic,<sup>24†</sup> E. Varenius,<sup>25†</sup> P. Santos,<sup>5†</sup> S. Seethapuram Sridhar,<sup>26,27†</sup> S. H. J. Wallström,<sup>25†</sup> J. Wiegert,<sup>25†</sup>

*Affiliations are listed at the end of the paper.*

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## ABSTRACT

We present ultraviolet (UV) observations of six nearby Type Ia supernovae (SNe Ia) obtained with the *Hubble Space Telescope*, three of which were also observed in the near-IR (NIR) with Wide-Field Camera 3. UV observations with the *Swift* satellite, as well as ground-based optical and near-infrared data provide complementary information. The combined data-set covers the wavelength range 0.2–2  $\mu\text{m}$ . By also including archival data of SN 2014J, we analyse a sample spanning observed colour excesses up to  $E(B-V) = 1.4$  mag. We study the wavelength dependent extinction of each individual SN and find a diversity of reddening laws when characterised by the total-to-selective extinction  $R_V$ . In particular, we note that for the two SNe with  $E(B-V) \gtrsim 1$  mag, for which the colour excess is dominated by dust extinction, we find  $R_V = 1.4 \pm 0.1$  and  $R_V = 2.8 \pm 0.1$ . Adding UV photometry reduces the uncertainty of fitted  $R_V$  by  $\sim 50\%$  allowing us to also measure  $R_V$  of individual low-extinction objects which point to a similar diversity, currently not accounted for in the analyses when SNe Ia are used for studying the expansion history of the universe.

**Key words:** ISM: dust, extinction – supernovae: general - circumstellar matter

## 1 INTRODUCTION

Studies of the cosmological expansion history using Type Ia supernovae (SNe Ia) have greatly improved our understanding of the Universe. While pioneering work lead to the discovery of the accelerated expansion from a few dozen SNe Ia (Riess et al. 1998; Perlmutter et al. 1999) the present samples of several hundred SNe Ia out to  $z \sim 1.5$  show to high-precision that Einstein’s Cosmological Constant provides an excellent fit for this phenomenon (e.g. Amanullah et al. 2010;

Conley et al. 2011; Sullivan et al. 2011; Suzuki et al. 2012; Betoule et al. 2014).

These findings have been possible due to the great accuracy of SNe Ia as distance indicators when observed in two or more broadband filters over several weeks, starting before light curve peak (see Goobar & Leibundgut 2011, for a recent review). The distance to each individual SN is obtained by fitting a SN Ia template to the photometric data (see e.g. Riess, Press & Kirshner 1996; Nugent, Kim & Perlmutter 2002; Jha, Riess & Kirshner 2007; Guy et al. 2007; Conley et al. 2008; Burns et al. 2011, and references therein) in order to obtain the *brightness* and *colour* at maximum together with the *light curve-shape*. The fitted properties are then typically combined using empirical linear relations to

\* rahman@fysik.su.se

† Participant of the Nordic Millimetre and Optical/NIR Astronomy Summer School 2012

form a distance-dependent quantity (Phillips 1993; Tripp 1998).

The colour correction is based on the observation that red SNe are fainter than bluer ones, which could originate from a combination of an intrinsic SN Ia colour law, and extinction by dust in the host galaxy. The colour-brightness relation is commonly determined by minimising the scatter around the best fit cosmological model, but there is no a priori reason why the two effects should follow the same colour-luminosity relation, nor can it be expected that host galaxy dust should have the same properties in different SN environments. In fact, even in the Milky Way, a range of extinction properties have been observed for different lines-of-sight.

When SNe Ia are used to measure distances, these uncertainties are handled by introducing a systematic error on the colour-luminosity relation, which makes an important contribution to the error budget when propagated to the derived cosmological parameters. Addressing the colour-brightness relation, and in particular breaking the degeneracy between the intrinsic colour component and dust extinction, is important for future SN Ia surveys to improve beyond the current cosmological constraints.

The properties and wavelength-dependent extinction of dust in the Milky Way has been carefully studied and is commonly characterised by the total,  $A_V$ , to selective,  $E(B-V) = A_B - A_V$ , extinction ratio as  $R_V = A_V / (A_B - A_V)$ . Lower values of  $R_V$  correspond to *steeper* extinction laws, since, for a given total extinction,  $A_V$ , these imply a larger reddening,  $E(B - V)$ .

The reddening of SNe Ia can be studied by comparing observed colours between reddened and similar un-reddened objects. Several such studies of individual SNe Ia (e.g. Elias-Rosa et al. 2006, 2008; Krisciunas et al. 2007; Folatelli et al. 2010; Amanullah et al. 2014, from hereon A14) in nearby galaxies, point to significantly lower values of  $R_V$  than has been observed in the Milky Way. Since the reddening of several of these SNe Ia is significantly higher than what can be expected from the Milky Way extinction along the line of sight, the observed colour excesses are likely dominated by extinction in the SN host galaxies. For low reddening SNe this may not be the case, but it is by nature difficult to measure  $R_V$  to high precision for individual SNe with low  $E(B - V)$ . However, the global reddening law for a sample of SNe can be measured, and several studies have obtained low  $R_V$  values (e.g. Astier et al. 2006; Nobili & Goobar 2008; Burns et al. 2014, from hereon B14) using this approach.

Chotard et al. (2011) and Scolnic et al. (2014) noted that globally derived reddening laws will be biased if the intrinsic SN Ia colour dispersions are not accounted for correctly, and Chotard et al. (2011) find  $R_V = 2.8 \pm 0.3$ , consistent with the Milky Way average, for their sample of 76 SNe Ia of which 73 had  $E(B - V) < 0.3$  mag. B14 also obtain  $R_V \approx 3$  when they constrain their analysis to only include low-extinction SNe Ia, while highly reddened objects appear to prefer lower values ( $R_V \approx 1.7$ ) which is in agreement with previous studies (Folatelli et al. 2010). This suggests that a single, global, reddening law derived from a set of SNe Ia could depend on the colour distribution of the sample.

The degeneracy between intrinsic SNe Ia colour variations and dust extinction, can also be approached by study-

ing the relation between SN colours and spectroscopic properties. Wang et al. (2009) found  $R_V = 1.6$  for SNe Ia with "high" photospheric velocities (HV,  $> 11,800$  km/s) within 5 days of maximum, while they obtained  $R_V = 2.4$  for objects with "normal" velocity (NV,  $< 11,800$  km/s). The expansion velocities were quantified by measuring the velocity of the Si II 6355 Å absorption feature. Foley & Kasen (2011) confirmed their findings but also conclude that when SNe with  $E(B - V) > 0.35$  mag are omitted, they obtain  $R_V \approx 2.5$  for the two separate subsamples, which also have different intrinsic colours. In other words, these findings also suggest that low  $R_V$  values are primarily associated with high-extinction, and these could dominate when a single global reddening law is derived from a sample of SNe Ia.

One natural explanation for low  $R_V$  values arises if SNe Ia are surrounded by circumstellar (CS) dust (Wang 2005; Goobar 2008). In this scenario, photons can scatter back into the line-of-sight which reduces the total extinction. Photons that scatter on the CS dust will also arrive later than photons that do not interact which will give rise to time-dependent reddening (Wang 2005; Amanullah & Goobar 2011; Brown et al. 2015). An observation will be a superposition of the SN spectrum at the given epoch and SN light from earlier epochs, but the scattering cross-section is wavelength-dependent so the fraction of delayed photons will also vary with wavelength. Further, depending on the geometry, scattered photons could give rise to a plateau for the late-time tail of the light curve and indications of such tails have been observed in  $B - V$  (Förster et al. 2013).

CS dust could also be heated by the SN, where the temperature will depend on the dust properties and the distance from the explosion. Heated dust has been detected in mid to far-infrared for a subset of peculiar SNe Ia (Fox et al. 2011; Taddia et al. 2012; Fox & Filippenko 2013; Silverman et al. 2013), but *Herschel* and *Spitzer* observations (Johansson, Amanullah & Goobar 2013; Johansson et al. 2014) of a handful of normal SNe Ia show no signs of heated CS dust (see also Maeda, Nozawa & Motohara 2014, for limits from Near-IR observations). For two of these SNe, 2006X and 2007le, together with SNe 1999cl and the SN 2002ic-like PTF11kx, time-varying Na I D absorption features have been observed in their spectra on time-scales comparable to the SN life-time which is consistent with the presence for CS material (Patat et al. 2007; Blondin et al. 2009; Simon et al. 2009; Dilday et al. 2012). Absorption from Na I has been used as a proxy for dust extinction (see e.g. Poznanski, Prochaska & Bloom 2012), although high-resolution spectra are required (Poznanski et al. 2011) and its validity for extinction along lines-of-sight of SNe Ia has been challenged (Phillips et al. 2013). Time-varying Na I D can be explained by photo-ionisation of neutral sodium (e.g. Borkowski, Blondin & Reynolds 2009) and subsequent recombination by the SN. In a recent study by Soker (2014) this explanation has been questioned and an alternative model of photon-induced desorption of sodium from dust in planetary nebulae remnants was proposed. It has also been suggested that time-variations may originate from patchy interstellar material along the line-of-sight (Patat et al. 2010). More recently, Graham et al. (2015b) claimed a detection of a time-varying K I absorption feature in 2014J, while the corresponding Na I D feature remained unchanged, which they argue is consistent with the presence of CS material at

a radius of  $r_{\text{dust}} \sim 10^{19}$  cm due to the different ionisation cross-sections.

An extensive study with multi-epoch high-resolution spectra of 14 SNe (Sternberg et al. 2014) have failed to reveal further examples of time varying Na I D. Nevertheless, a statistically significant preponderance of Na I D features which are blue shifted with respect to the local velocity have been observed in high- and mid-resolution spectra (Sternberg et al. 2011; Maguire et al. 2013, respectively), suggesting that there is outflowing material from the SN progenitor system.

In this work we seek to expand our knowledge of extragalactic extinction in the line-of-sight and SN Ia reddening by making the widest wavelength study of a sample of SNe Ia to date. We present UV-to-NIR light curves for six SNe Ia based on data mainly obtained from The Nordic Optical Telescope (NOT) and *The Hubble Space Telescope* (*HST*). The observations and data reduction is presented in §2 while the light curves and colours are presented in §3. The method for deriving individual reddening laws for each SN is described in §4, and after adding SN 2014J from A14 to the sample we present the results in §4.3. We search for CS dust in §5 and all the results are discussed in §6. We conclude and summarise in §7.

All seven SNe<sup>1</sup> are listed in Table 1 and shown together with their host galaxies in Figure 1. They are also briefly summarised in Appendix A.

## 2 OBSERVATIONS AND DATA REDUCTION

### 2.1 Hubble Space Telescope

All previously unpublished *HST* observations discussed in this work are listed in Table 2. Each SN was observed at two different epochs (four for SN 2012et) with the Wide-Field Camera 3 (WFC3) during the *HST* Cycle 19 under programme GO-12582 (PI: Goobar). All SNe were imaged with the *WFC3/UVIS* channel through the passbands F225W, F275W and F336W. In addition to this, NIR imaging in the *HST* passbands F125W and F160W was obtained for the SNe 2012cg, 2012cu and 2012et with the *WFC3/IR* channel. Examples of both the *WFC3/UVIS* and *WFC3/IR* observations are shown in Figure 2.

*WFC3/UVIS* consists of two  $4096 \times 2051$  e2V CCD detectors with a plate scale of  $0.04''/\text{pixel}$ , while the Teledyne HgCdTe infrared detector used in the *WFC3/IR* channel has a pixel scale of  $0.13''/\text{pixel}$ . For each observation, only the sub-arrays (apertures), listed in Table 2, of the detectors were read-out. All WFC3 data were reduced using the standard STScI reduction pipeline and calibrated using CALWF3.

The *WFC3/UVIS* CCDs, like all *HST* CCDs, are plagued with charge transfer inefficiencies (CTI), i.e. degradations of the detector performance over time due to damage in the silicon lattice from cosmic rays. The CTI can be reverse corrected at the pixel level using `wfc3uv.ctereverse`<sup>2</sup>

<sup>1</sup> All data and figures presented in this work are available at <http://snova.fysik.su.se/dust/>.

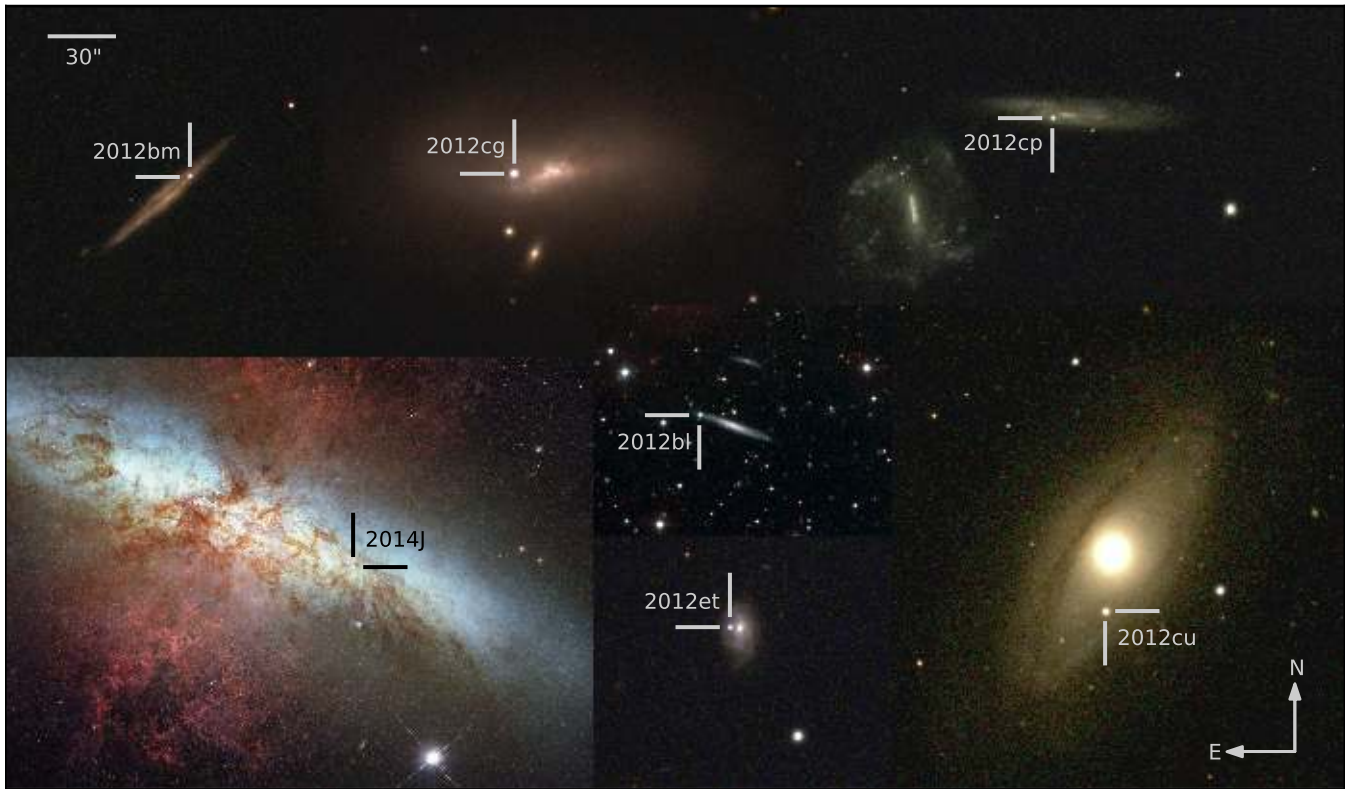
<sup>2</sup> [http://www.stsci.edu/hst/wfc3/tools/cte\\_tools](http://www.stsci.edu/hst/wfc3/tools/cte_tools)

Civil date	MJD	Aperture	SN
2012-04-09	56026.4	UVIS1-2K2A-SUB	2012bm
2012-04-09	56026.5	STIS 52X0.2 G230LB	2012bm
2012-04-13	56030.6	UVIS1-2K2A-SUB	2012bm
2012-04-13	56030.7	STIS 52X0.2 G230LB	2012bm
2012-04-16	56033.1	STIS 52X0.2 G230LB	2012bl
2012-04-16	56038.2	UVIS1-2K2A-SUB	2012bl
2012-04-20	56037.4	STIS 52X0.2 G230LB	2012bl
2012-04-21	56038.2	UVIS1-2K2A-SUB	2012bl
2012-06-04	56082.4	UVIS2-M1K1C-SUB	2012cg
2012-06-04	56082.4	IRSUB64	2012cg
2012-06-04	56082.5	STIS 52X0.2 G230LB	2012cg
2012-06-04	56082.5	STIS 52X0.1 G430L	2012cg
2012-06-18	56096.4	STIS 52X0.2 G230LB	2012cg
2012-06-18	56096.4	STIS 52X0.1 G430L	2012cg
2012-06-19	56097.5	UVIS2-M1K1C-SUB	2012cg
2012-06-19	56097.5	IRSUB64	2012cg
2012-06-04	56082.6	UVIS1-2K2A-SUB	2012cp
2012-06-04	56082.7	STIS 52X0.2 G230LB	2012cp
2012-06-16	56094.6	UVIS1-2K2A-SUB	2012cp
2012-06-17	56095.6	STIS 52X0.2 G230LB	2012cp
2012-07-02	56110.8	UVIS2-M1K1C-SUB	2012cu
2012-07-02	56110.8	IRSUB512	2012cu
2012-07-07	56115.8	UVIS2-M1K1C-SUB	2012cu
2012-07-07	56115.8	IRSUB512	2012cu
2012-10-01	56201.5	UVIS2-M1K1C-SUB	2012et
2012-10-01	56201.5	IR	2012et
2012-10-01	56201.8	STIS 52X0.2 G230LB	2012et
2012-10-01	56201.8	STIS 52X0.2 G430L	2012et
2012-10-05	56205.6	UVIS2-M1K1C-SUB	2012et
2012-10-05	56205.6	IR	2012et
2012-10-09	56209.8	UVIS2-M1K1C-SUB	2012et
2012-10-09	56209.8	IR	2012et
2012-10-13	56213.9	UVIS2-M1K1C-SUB	2012et
2012-10-13	56213.9	IR	2012et

**Table 2.** Wide-Field Camera-3 and STIS observations together with the used apertures, where UVIS1 and UVIS2 are the two chips of the *WFC3/UVIS* channel. The observations with the *WFC3/IR* and *WFC3/UVIS* channels were obtained through the F125W, F160W and F225W, F275W, F336W broadband passbands respectively.

for all frames except for images obtained with the UVIS2-M1K1C-SUB sub-array. For this aperture we instead used the recipe suggested by Bourque & Anderson (2014). The individual flat-fielded *WFC3/UVIS* and *WFC3/IR* images were then resampled and combined using `astrodrizzle` (Fruchter & et al. 2010).

Aperture photometry was carried out on the combined images using an aperture radius of  $0.4''$ , although in some cases a smaller aperture of radius  $0.16''$  had to be used for the *WFC3/UVIS* data. This was the case for the low signal-to-noise measurements of SN 2012bm, observations of SN 2012cg that were plagued by nearby coincidence cosmic rays, and all epochs of SN 2012et, where a nearby object contaminated the larger aperture (see the mid-right panel of Figure 2). The smaller radius measurements had to be aperture corrected to the  $0.4''$  radius for which STScI pro-



**Figure 1.** The seven SNe analysed in this work. The patches for SNe 2012bm, 2012cg, 2012cp, 2012cu and 2012et were obtained from *BVR* images from NOT, while the images for SNe 2012bl and 2014J are obtained from the Swope telescope and *HST/WFC3* (Program DD-13621; PI Goobar) images respectively.

SN	$\alpha_{\text{SN}}$ (2000)	$\delta_{\text{SN}}$ (2000)	$r_{\perp}$ (kpc)	Host galaxy	$\alpha_{\text{h}}$ (2000)	$\delta_{\text{h}}$ (2000)	$v_{\text{h}}$ (km/s)	Distance (Mpc)	MW $A_V$ (mag)
2012bl	20:23:55.28	-48:21:17.3	15.6(0.2)	ESO 234-19	20:23:51.0	-48:21:32	5608	71(1) <sup>a</sup>	0.098
2012bm	13:05:45.66	+46:27:52.9	7.5(0.3)	UGC 8189	13:05:46.6	+46:27:42	7436	103(8) <sup>b</sup>	0.033
2012cg	12:27:12.83	+09:25:13.2	1.3(0.2)	NGC 4424	12:27:11.6	+09:25:14	437	15(2) <sup>c</sup>	0.057
2012cp	13:47:01.79	+33:53:35.0	1.9(0.1)	UGC 8713	13:47:01.2	+33:53:37	4956	57(3) <sup>d</sup>	0.058
2012cu	12:53:29.35	+02:09:39.0	5.4(1.2)	NGC 4772	12:53:29.1	+02:10:06	1040	41(9) <sup>e</sup>	0.074
2012et	23:42:38.82	+27:05:31.5	2.7(0.2)	MCG +04-55-47	23:42:38.4	+27:05:31	7483	105(7) <sup>a</sup>	0.221
2014J	09:55:42.11	+69:40:25.9	1.0(0.3)	M82	09:55:52.7	+69:40:46	203	3.5(0.3) <sup>f</sup>	0.189 <sup>f</sup>

<sup>a</sup> Parnovsky & Parnowski (2010); <sup>b</sup> Calculated based on the redshift with  $H_0 = 73 \pm 5$  km/s/Mpc; <sup>c</sup> Cortés, Kenney & Hardy (2008); <sup>d</sup> Springob et al. (2009); <sup>e</sup> Tully et al. (2009); <sup>f</sup> Dalcanton et al. (2009)

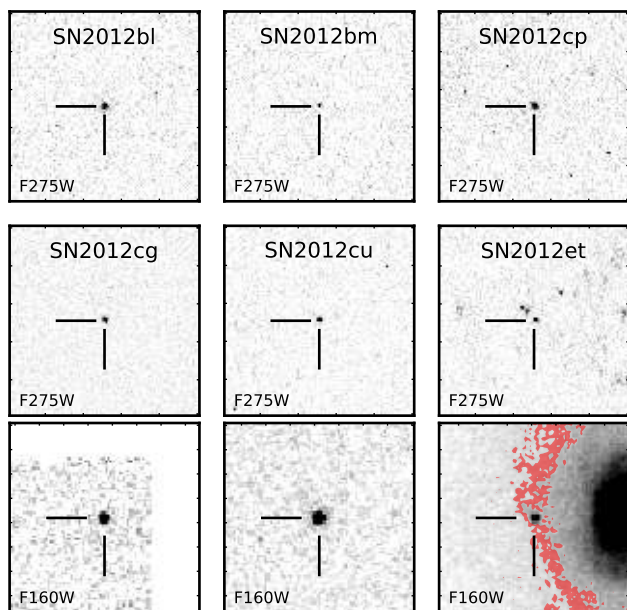
**Table 1.** SN coordinates are quoted from the discovery telegrams. All host galaxy data were obtained from the Nasa Extragalactic Database (NED), unless otherwise specified, where  $v_{\text{h}}$  is the measured recession velocity, and the Milky Way extinctions are from the Schlafly & Finkbeiner (2011) calibration of the Schlegel, Finkbeiner & Davis (1998) infrared-based dust maps. The projected distances from the host galaxy nuclei,  $r_{\perp}$ , were calculated based on the host galaxy distances and the SN offsets specified in Appendix A.

vides zeropoints<sup>3</sup> for WFC3. Since there were no bright stars present in the fields we used the general enclosed energy tables provided by STScI for this purpose. To account for the time and wavelength dependence of the *WFC3/UVIS* aperture corrections we adopt an additional uncertainty of 0.05 mag for all measurements using the 4 pixel radius, which was derived by comparing the enclosed energy tables

with photometry from different aperture radii of the high signal-to-noise data.

Only the SN 2012et NIR measurements suffered from host galaxy contamination. The host contribution was estimated by placing apertures of the same radius used for the SN photometry along the isophot of the galaxy that intersects with the position of SN 2012et as illustrated in the lower right panel of Figure 2. The background was then estimated as the median of these measurements, and the root-mean-square of the measurements was added in quadrature to the photometric uncertainty. All the *WFC3* photometry

<sup>3</sup> We used the zeropoints from March 6th, 2012.



**Figure 2.** *HST/WFC3* patches of the six SNe observed during Cycle 19. The patch sizes are  $10'' \times 10''$  and in the upper row the first epochs obtained through the F275W filter is shown for each of the three SNe 2012bl, 2012bm and 2012cp. In the two lower rows, the corresponding patches are shown for the SNe 2012cg, 2012cu and 2012et, together with the NIR observations in the F160W filter that also were obtained for these objects. The marked (red) region in the lower right panel show the isophot from which the host galaxy background of SN 2012et was estimated as explained in the text.

is presented in Table 3 and shown in Figure 3. Here we also present updated photometry of SN 2014J using the method described here which is consistent with the results presented in A14.

For most *HST* visits we also obtained long-slit spectroscopy with the Space Telescope Imaging Spectrograph (STIS) where the  $1024 \times 1024$  pixel SiTe CCD detector was used with the G230LB grating covering the wavelength range 1600–3100 Å, and for two SNe (2012cg, 2012et) G430L, 2900–5700 Å. The data were reduced using the *calstis* pipeline which is part of the *STSDAS* package. The pipeline was ran for all the spectra up to the point where the calibrated 2D spectra were created. Only the spectra of SN 2012cg and the red G430L spectrum of SN 2012et contained any significant signal and were extracted using the *calstis* *x1d* routine with a 4 pixel aperture.

## 2.2 *Swift/UVOT*

UV photometry was also obtained with the Ultraviolet/Optical Telescope (*UVOT*, Roming et al. 2005) on the *Swift* spacecraft (Gehrels et al. 2004) for the SNe 2012bl, 2012cg, 2012cp and 2012cu in the *uvw1*, *uvw2* and *uvm2* filters. However, the *uvw1* and *uvw2* filters are not well suited for extinction studies due to the significant “red-tails” of these filters. Milne et al. (2010) used the UV spectra of SN 1992A (Kirshner et al. 1993) and estimated that 52% and 44% of the light in *uvw1* and *uvw2* respectively, originate from wavelengths redder than 3000 Å. Brown et al.

(2015) show how the observed photons and the corresponding effective wavelengths shift dramatically as reddening increases.

The *uvm2* filter, on the other hand, is much better constrained (1% of the light comes from  $> 3000$  Å) and is used for the extinction studies in this work. The *uvm2* magnitudes were measured using the pipeline from the *Swift* Optical/Ultraviolet Supernova Archive (SOUSA, Brown et al. 2014), including correction for the time-dependent sensitivity and revised zeropoints from Breeveld et al. (2011). The result is presented in Table 3 and shown in Figure 3. Only the SNe 2012bl and 2012cg showed a significant SN flux in the *uvm2* band after subtraction of the underlying host galaxy flux. The flux measured for SN 2012cp was consistent with the host galaxy brightness and no signal was measured for SN 2012cu.

## 2.3 Ground based observations

The ground based spectroscopic observations are listed in Table 4 while the photometric observations are summarised together with the measured magnitudes in Table 3 and shown in Figure 3. All SNe except for SN 2012bl were observed with the 2.56 metre Nordic Optical Telescope (NOT) under programmes 45-009 and 46-018 (PI: Amanullah). Imaging and spectroscopy was obtained with the  $6.4' \times 6.4'$  Andalucia Faint Object Spectrograph and Camera (ALFOSC) using the filter set  $U$  (#7)<sup>4</sup>,  $B$  (#74),  $V$  (#75),  $R$  (#76),  $i$  (#12) and the  $R = 360$  grism (#4)<sup>5</sup>. The data were reduced using standard IRAF routines and the QUBA pipeline (Valenti et al. 2011). The photometry was measured by first fitting the point-spread-function (PSF) using DAOPHOT (Stetson 1987) to stars in the fields. The calibrated magnitudes were obtained either from Landolt fields (Landolt 1992) observed during the night, or, when these were not available, by comparing the fluxes to stars in the fields that were then calibrated against Landolt fields for the photometric nights.

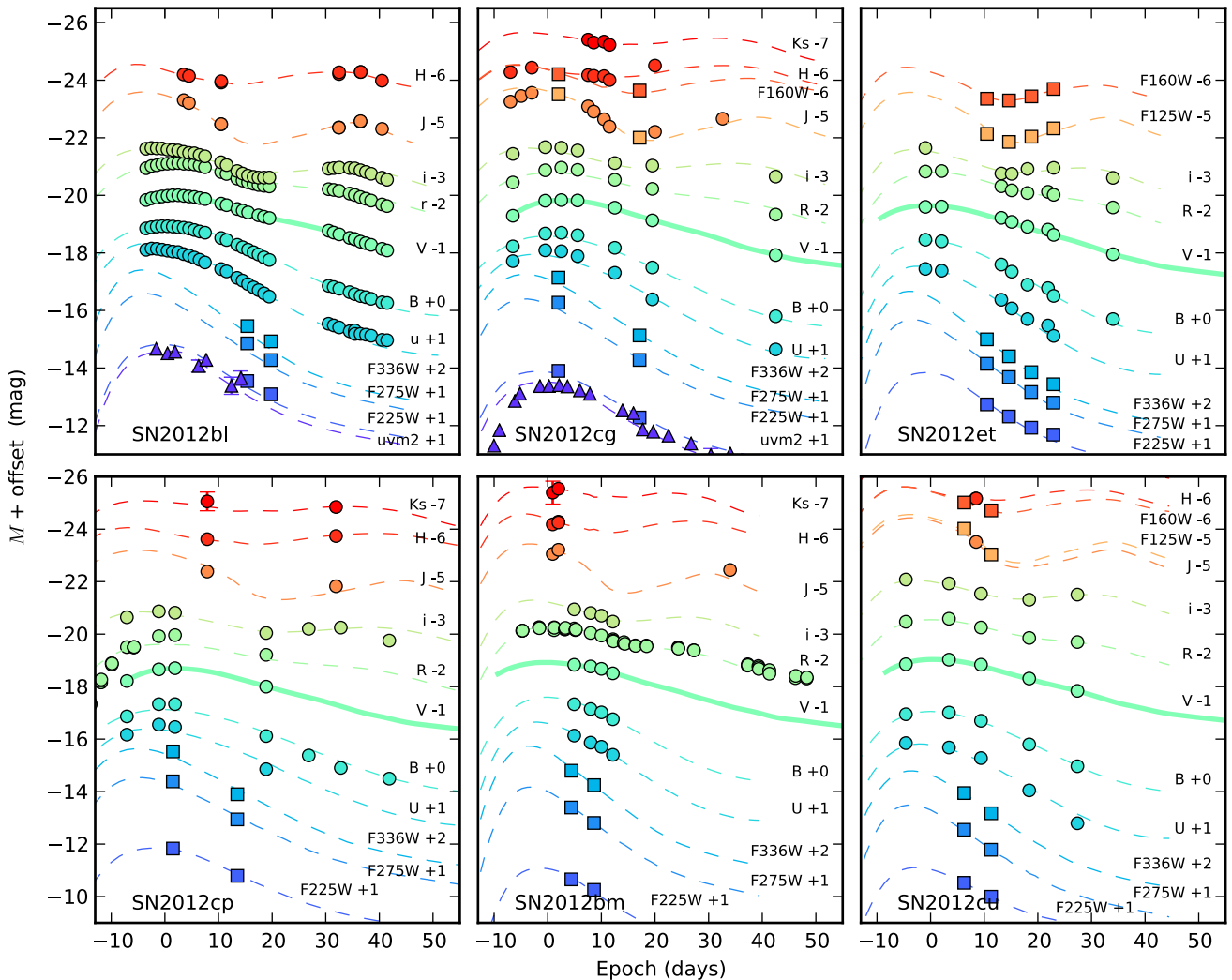
With the NOT we also obtained high-resolution spectroscopy of SN 2012cg using the Fibre-fed Echelle Spectrograph (FIES, Telting et al. 2014) in its high-resolution mode,  $R = 67000$ . Simultaneous wavelength reference (Thorium-Argon) spectra were obtained and the data were reduced using the software *FIESTool* which is provided by the observatory.

NIR observations were carried out for SNe 2012bm, 2012cg and 2012cp with the  $4' \times 4'$  NOTCam instrument in the  $J$ ,  $H$  and  $K_s$  bands under the programmes 45-007 (PI: Kankare) and 46-020 (PI: Mattila). We used the wide field imaging option of the  $1024 \times 1024$  pixel HgCdTe NOTCam detector with a plate scale of  $0''.234/\text{pixel}$ . The observations were carried out using either 5 or 9-point dithering patterns. For the SNe with extended host galaxies, beam-switching was used to guarantee a successful sky subtraction. We used the NOTCam Quick-Look reduction package based on IRAF. Bad pixels (which includes two dead columns) were masked and we used master differential skyflats which

<sup>4</sup> NOT filter ID

<sup>5</sup> NOT grism ID





**Figure 3.** Observed light curves for the SN sample. Our analysis is based on optical *UBVRI* and NIR *JHK* photometry from ground based observatories (circles), UV (F225W,F275W,F336W) and NIR (F125W,F160W) photometry from the *HST* (squares). For the SNe 2012bl and 2012cg we also use *Swift/UVOT* photometry (triangles). The data are shown in their instrumental system but have been shifted using the distance moduli based on the distances in Table 1 and offset for clarity. The thick, solid, lines show the SNooPy *V*-models used to calculate colours as described in the text, while the dashed lines show synthesised light curves of SN2011fe to guide the eye. When a SN has observations in a similar filter from multiple instruments only one of the light curves is shown. The epoch is the observed date with respect to the fitted day of maximum in the *B*-band from Table 5.

were obtained from bright and faint skyflats. Further, the images were corrected for geometric distortions<sup>6</sup> before the individual images were aligned and coadded.

SN 2012cg was also observed in *JHKs* with the  $4.3' \times 4.3'$  CAIN instrument on the 1.52 metre Carlos Sanchez Telescope at Observatorio del Teide on Tenerife. CAIN III is a  $256 \times 256$  pixel HgCdTe NIR detector, where we used the wide field option with a plate scale of  $1.0''/\text{pixel}$ . The data were reduced using a dedicated IRAF package provided by J. Pullido and A. Barrena.

The two SNe 2012bm and 2012cp were also observed by the *Palomar Transient Factory* with the Palomar 48-inch telescope (Rau et al. 2009) in the Mould *R*. We use

their photometry which was obtained by first determining the PSF prior to subtractions, and then carrying out PSF photometry on the subtracted frames (Firth et al. 2015). The calibration of this data has been described in Ofek et al. (2012).

Photometric follow-up of SN2012bl was carried out by the *Carnegie Supernova Project* using the Swope telescope for the optical and the du Pont telescope for the NIR. The reduction and photometry have been described in detail in Hamuy et al. (2006) and Contreras et al. (2010). PSF photometry was performed with respect to a local sequence of standard stars calibrated to the Landolt (1992) and Smith et al. (2002) standard fields for the optical observations and the Persson et al. (1998) for the NIR.

Further, we here also present spectroscopic observations of SNe 2011fe, 2012cg and 2014J obtained using the ISIS

<sup>6</sup> Magnus Gälfalk, private communication.

MJD (days)	Phase (days)	Filter	$X$ (mag)	$A_X^{\text{MW}}$ (mag)	$K_X$ (mag)	Match	$V$ (mag)	$A_V^{\text{MW}}$ (mag)	$K_V$ (mag)	$(V - X)_0$ (mag)
SN2012cg										
56082.4	0.2	WFC3 F336W	11.74(0.06)	0.09	0.01	M	12.04(0.04)	0.06	0.00	0.61
56097.5	14.0	WFC3 F336W	13.75(0.01)	0.09	0.02	M	12.53(0.07)	0.06	0.00	-0.90
56082.4	0.2	WFC3 F125W	12.37(0.01)	0.01	0.05	M	12.04(0.04)	0.06	0.00	-0.55
56097.5	14.1	WFC3 F125W	13.88(0.03)	0.02	-0.07	M	12.53(0.07)	0.06	0.00	-1.91
56082.4	0.2	WFC3 F225W	16.38(0.05)	0.13	0.02	M	12.04(0.04)	0.06	0.00	-2.55
56097.5	14.0	WFC3 F225W	17.60(0.08)	0.13	0.02	M	12.53(0.07)	0.06	0.00	-3.68
56082.4	0.2	WFC3 F160W	12.67(0.02)	0.01	0.01	M	12.04(0.04)	0.06	0.00	-0.88
56097.5	14.1	WFC3 F160W	13.24(0.03)	0.01	0.55	M	12.53(0.07)	0.06	0.00	-0.69
...										

**Table 3.** The photometry of all SNe. All magnitudes are in the natural Vega system. The rest-frame magnitude and colour can be obtained from Columns 4–6 using Equation (1). Columns 5 and 9 are the Galactic extinctions and Columns 6 and 10 are  $K_X$  and  $K_V$  corrections for the two bands, respectively. All corrections were calculated after the SN 2011fe template had been reddened with the best fitted F99 law, shown in Table 5, for each SN. The  $V$  magnitude is only shown for data points used in the colour analysis, i.e. with phases between  $-10$  and  $+35$  days. Column 2 show the effective light curve-width-corrected phase, while Column 7 specifies whether the  $V$  magnitude was measured for the same epoch (D) or if it was calculated using the SNOOPy model (M). The corresponding intrinsic colour for the SN 2011fe template is shown when available in Column 11. The *Swift* magnitudes of SN 2014J are from Brown et al. (2015). (This table is available in its entirety in a machine-readable form in the online journal. A portion is shown here for guidance regarding its form and content.)

Civil date	MJD	Epoch (days)	Instrument	$\lambda$ (Å)	Exp. (s)	SN
2012-05-26	56073.8	+260	WHT/ISIS	3500–9500	600	2011fe
2012-03-30	56016.3	-2	DuPont/WFCCD/WF4K-1	3600–9200	700	2012bl
2012-03-31	56017.4	-1	DuPont/WFCCD/WF4K-1	3600–9200	700	2012bl
2012-05-01	56048.3	+29	DuPont/WFCCD/WF4K-1	3600–9200	900	2012bl
2012-03-28	56014.0	-4	ASIAGO/AFOSC	3500–8200	2700	2012bm
2012-04-09	56027.0	+9	NOT/ALFOSC	3200–9000	2400	2012bm
2012-04-16	56034.1	+16	NOT/ALFOSC	3200–9000	2400	2012bm
2012-05-24	56071.9	-9	NOT/ALFOSC	3700–7200	5477	2012cg
2012-05-26	56073.9	-7	NOT/FIES	3640–7360	1800	2012cg
2012-05-28	56075.9	-5	NOT/ALFOSC	3700–7200	6588	2012cg
2012-06-02	56080.9	+1	NOT/FIES	3640–7360	900	2012cg
2012-06-04	56082.9	+3	NOT/ALFOSC	3200–9000	300	2012cg
2012-06-14	56092.9	+13	NOT/FIES	3640–7360	1200	2012cg
2013-03-09	56360.4	+280	Keck/LRIS	3075–10300	600	2012cg
2012-05-26	56074.0	-7	NOT/ALFOSC	3200–9000	2700	2012cp
2012-06-04	56083.0	+2	NOT/ALFOSC	3200–9000	1800	2012cp
2012-07-05	56113.9	+9	NOT/ALFOSC	3200–9000	1200	2012cu
2012-09-13	56183.8	-7	ASIAGO/AFOSC	3500–8200	2700	2012et
2012-09-22	56193.1	+2	NOT/ALFOSC	3200–9000	600	2012et
2014-11-09	56970.0	+282	APO/DIS	3300–9800	750	2014J

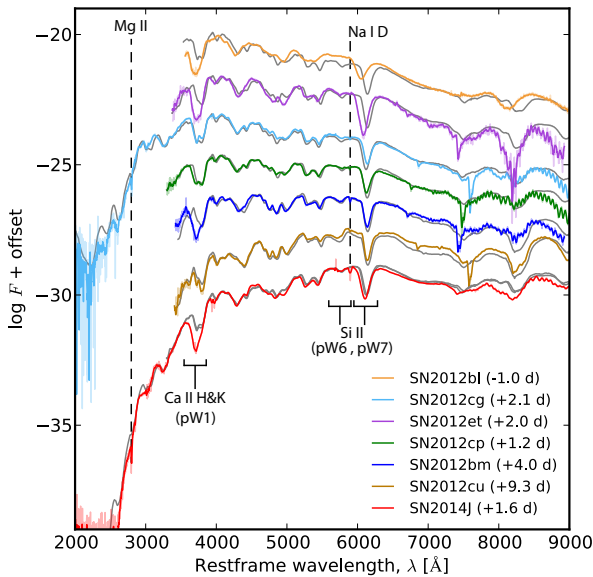
**Table 4.** Ground based spectroscopic observations. The Asiago spectra were obtained through the Asiago Transient Classification Program (Tomasella et al. 2014). The epochs are shown with respect to the date of  $B$ -maximum shown in Table 5.

instrument at the William Herschel Telescope, the LRIS instrument at Keck and the DIS instrument at the Apache Point Observatory. These data were reduced in a standard manner.

### 3 SN SPECTRA AND COLOURS

The spectra for the seven SNe around maximum are shown in Figure 4, where we also added the classification spectrum of SNe 2012bm and 2012et (Tomasella et al. 2014) to the set of spectroscopic observations already described above. For SN 2012cg, both the *HST/STIS* G250LB and G430L spectra, and the ALFOSC spectrum, that was ob-





**Figure 4.** Spectra near maximum, when available, for the SNe in our sample. The grey spectra are for SN 2011fe at similar epochs (Pereira et al. 2013; Mazzali et al. 2014) reddened by the extinction laws that we derive in § 4. For SNe 2012cg and 2014J we have combined the optical spectra with *STIS* spectra (the SN 2014J spectrum was taken from Foley et al. 2014). Spectroscopic features discussed in § 5, § 6.1 and § 6.4 have been marked.

tained the same day, are shown. We used SNID (Blondin & Tonry 2007) to type the spectra and all SNe could be classified as “normal”, although SN 2012et also provided acceptable matches to SN 1999aa. For each spectrum in Figure 4 we also show the spectrum of the normal SN 2011fe for the matching epoch (Mazzali et al. 2014) for comparison.

All photometric data are shown in Figure 3 and for each measurement in each filter,  $X$ , where  $X \in \{um2, F225W, \dots, Ks\}$ , we calculated the colour  $X - V$ . When  $X$  and  $V$ -band measurements are not available for the same dates, and the colour could not be obtained directly, a  $V$ -band model was used to calculate the colours. When applicable, we used a smoothed spline model but for the SNe 2012bm and 2012cp the  $V$ -band data were too sparsely sampled to allow reliable spline fits. For these objects we instead fitted the SN Ia template from Hsiao et al. (2007, H07 from hereon) to the observed data. Both the spline and template-based models were obtained using the *SNooPy* light curve fitter (Burns et al. 2011) and are shown as thick lines in Figure 3. Further, for all model-based colours we add 0.05 mag in quadrature to uncertainties in order to account for the inaccuracy of the models.

All colours were corrected for Galactic extinction,  $A_V^{\text{MW}}$ , using the extinction law by Cardelli, Clayton & Mathis (1989, CCM from hereon) with  $R_V = 3.1$  and the  $A_V^{\text{MW}}$  values shown in the last column of Table 1. The values are from the Schlafly & Finkbeiner (2011) recalibration of the Schlegel, Finkbeiner & Davis (1998) dust maps, except for SN 2014J where we use a value obtained from studying neighbouring regions (Dalcanton et al. 2009).

The colours were also  $K$  corrected (e.g. Nugent, Kim & Perlmutter 2002) and  $S$  corrected (e.g. Suntzeff 2000; Stritzinger et al. 2002; Krisciunas et al. 2003) to a common rest-frame filter system that will be used for the remaining analysis throughout the paper. The analysis was carried out using the *WFC3/UVIS* and *Swift/UVOT* filters for the UV, the ALFOSC filter set for the optical and the NOTCam filters for the NIR, since the bulk of the data were obtained using these filters. The combined  $K$  and  $S$  corrections ( $K_X$  from now on) are calculated synthetically using the SED of SN 2011fe as described in § 4.1 and the filter transmissions provided by the different observatories.

To summarise, all colours were obtained as

$$X - V = (m_X - A_X^{\text{MW}} - K_X) - (\mathcal{M}_V - A_V^{\text{MW}} - K_V), \quad (1)$$

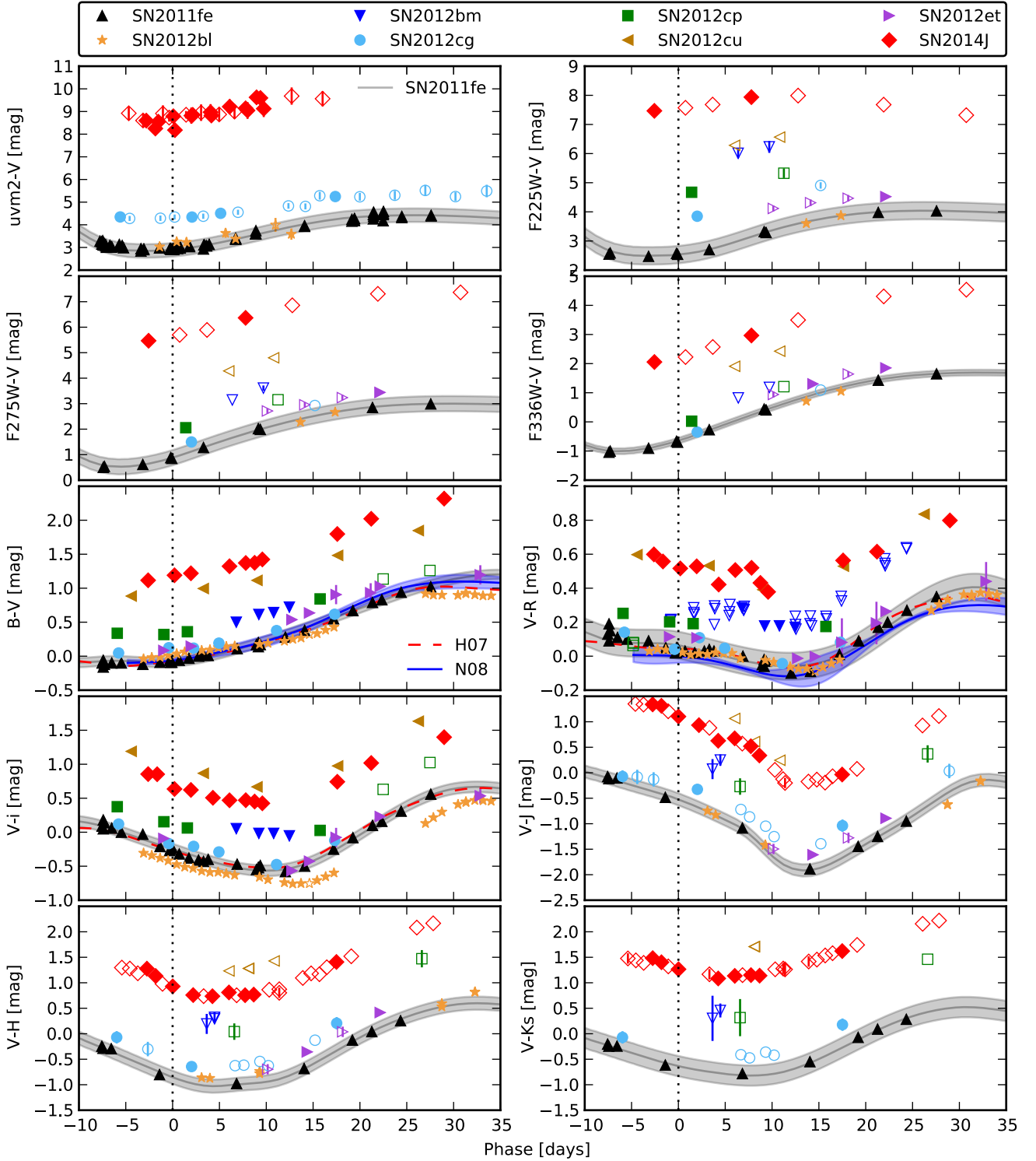
where  $m_X$  are the measured instrumental magnitudes in filter  $X$  and  $\mathcal{M}_V$  is the  $V$  magnitude for the same date either measured directly, or obtained from a spline or template model. All values are presented in Table 3 where we have also added the *Swift* photometry of SN 2014J from Brown et al. (2015) to the measurements described in § 2. As will be further discussed in § 4, not only the observed colours, but also,  $A^{\text{MW}}$  and  $K_X$  depend on the reddening of each SN and were therefore obtained iteratively during the extinction law fitting. The values shown in Table 3 were obtained after the SN 2011fe SED had been reddened with the Fitzpatrick (1999, hereafter F99) law using the parameters shown in Table 5 for each SN.

A subset of the colours are also shown in Figure 5 where the colours obtained using measurements and the  $V$ -band model are shown with filled and open symbols respectively.

SNe Ia are intrinsically bright in the rest-frame  $B$ -band and this filter is commonly used to measure SN Ia distances. SN Ia light curves are typically quantified by a few parameters such as the brightness and time of maximum,  $t_B^{\text{max}}$ , for e.g. the  $B$ -band, and the light curve shape. The light curve shape can be quantified using different methods, where perhaps the most straight-forward is to measure the brightness decline for the first 15 days past  $t_B^{\text{max}}$ . Phillips (1993) showed that fast declining SNe Ia are typically intrinsically fainter than slow declining objects, a discovery that lay the foundation for using SNe Ia as distance indicators for precision cosmology. An alternative approach to quantify the light curve shape is to introduce a *stretch* parameter (Perlmutter et al. 1997),  $s_B$ , which can be defined as the value needed to match the time-evolution of an observed  $B$  light curve to a standard SN Ia template.

The time-evolution of SN Ia  $X - V$  colours also depends on light curve shape. For example, Nobili & Goobar (2008, from hereon N08) present colour-light curve shape relations in the optical while B14 show that the SNe Ia colours can be standardised over a wide range of decline rates by introducing *colour-stretch*,  $s_{BV} = t_{BV}/30$  days, where  $t_{BV}$  is the time between  $t_B^{\text{max}}$  and the maximum of the  $B - V$  colour. However, this approach requires  $B - V$  coverage up to  $t \sim t_B^{\text{max}} + 40$  in order to accurately determine  $s_{BV}$ .

We here standardise the measured colours using  $s_B$ . All colours will be studied and compared as a function of *phase*,  $p$ , defined as  $p = (t - t_B^{\text{max}})/s_B$ , where  $t$  are the observing dates. The phases for the colours shown in Figure 5 have been determined by using the values of  $t_B^{\text{max}}$  and  $s_B$  from



**Figure 5.** A selection of measured colours from UV to NIR together with literature SNe 2011by and 2011fe. The phases,  $p = (t - t_B^{\max})/s_B$ , were obtained using the values of  $t_B^{\max}$  and  $s_B$  from Table 5. The colours from the H07 spectral template (dashed red) and Nobili & Goobar (2008) (N08, solid blue with dispersion region) are also shown together with the derived colour model of SN 2011fe (solid grey) with the adopted dispersions (grey region). Filled symbols show colours where a  $V$ -band measurement was obtained for the same date, while a  $V$ -band model was used to obtain the measurements that are shown with open symbols. Errorbars are only plotted for data points where the uncertainty exceeds 0.1 mag. See text for details.

Table 5, which are fitted simultaneously with the reddening law as described in § 4.

#### 4 FITTING REDDENING LAWS

We will use a similar method to A14 to fit different parametrised reddening laws to the measured colour excesses,

$$E(X - V) = (X - V) - (X - V)_0, \quad (2)$$

where  $(X - V)_0$  is the assumed  $X - V$  colour for the unreddened objects.

In this work we will study different extinction laws and how well they describe the observed SN Ia reddening. Each extinction law,  $A(\lambda; \bar{a})$ , depends on a set of parameters  $\bar{a}$  (e.g.  $E(B - V)$  and  $R_V$ ), that will be fitted to the observed colour excesses,  $E(X - V)_p$ , by minimising

$$\chi^2 = \sum_X \sum_p \frac{[E(X - V)_p - (A_{X_p} - A_{V_p})]^2}{\sigma_{X_p V_p}^2}. \quad (3)$$

Here  $A_{X_p}$  and  $A_{V_p}$  are the predicted extinctions in the  $X$  and  $V$  filters for phase  $p$  and can be calculated as

$$A_{X_p} = -2.5 \log_{10} \left( \frac{\int T_X(\lambda) \cdot 10^{-0.4 \cdot A(\lambda; \bar{a})} \cdot S_0(\lambda; p) \lambda d\lambda}{\int T_X(\lambda) S_0(\lambda; p) \lambda d\lambda} \right), \quad (4)$$

if the effective filter transmission,  $T_X(\lambda)$ , and the spectral energy distribution (SED),  $S_0(\lambda; p)$ , of the unreddened source are assumed to be known. Further,  $\sigma_{X_p V_p}^2$  are the uncertainties added in quadrature. This includes the measurement errors, but will in most cases be dominated by the intrinsic colour uncertainties discussed in § 4.1.

In addition to the extinction law parameters  $\bar{a}$ , we will also minimise Equation (3) with respect to the parameters  $t_B^{\max}$  and  $s_B$  that together determine the phase,  $p$ . The *observed* values of these parameters are extinction dependent (see e.g. Leibundgut 1988; Phillips 1993; Nugent, Kim & Perlmutter 2002, A14) since broadband measurement of an object suffering from extinction will effectively probe redder wavelengths than the observation of the same unreddened object would have. The SN Ia light curve decline rate varies with wavelength, and slower decline rates are in general expected to be observed for reddened objects.

Similarly, the Galactic extinction,  $A_X^{\text{MW}}$ , as well as the  $K_X$  corrections, are also properties measured through broadband filters and therefore also depend on the SED of the source, as seen from e.g. Equation (4). Since the observed SED in the Milky Way is affected by the extinction in the SN host galaxy, these properties will also depend on the reddening law. The effect will be particularly significant in the UV, where the wavelength dependence for the extinction laws is steep. For example, the Galactic extinction in the *Swift*/*uvm2* band for the line of sight to SN 2014J would have been  $A_X^{\text{MW}} \approx 0.3$  for an unreddened source while it is  $A_X^{\text{MW}} \approx 0.1$  for SN 2014J.

We take this effect into account by fitting each reddening law,  $A(\lambda; \bar{a})$ , iteratively and update  $A_X^{\text{MW}}$  and  $K_X$  using the fit  $\bar{a}$  values in each iteration. That is, after each iteration, we re-calculate the observed colour and colour excesses using Equations (1) and (2) before re-fitting the reddening law using Equation (3). The procedure is repeated until the

change in the fitted parameters is less than 1% between iterations.

As mentioned above we use the average value of  $R_V = 3.1$  when correcting for the Galactic extinction. For the majority of the SNe the Galactic extinction is negligible compared to the host reddening, with the exception of SN 2012et. For this SN we also tried carrying out the fits using the extreme values of  $R_V = 2.2$  and  $R_V = 5.8$  observed in the Milky Way (e.g. CCM, Fitzpatrick 1999) and concluded that this could impact the fit values with up to one statistical standard deviation.

#### 4.1 SN Ia intrinsic colours and SED

In order to minimise Equation (3), we need to make assumptions of the intrinsic colours,  $(X - V)_0$ , and SED,  $S_0(\lambda; p)$  of the unreddened objects. The reddening laws for SNe Ia are typically derived by either comparing them to individual objects that show similarities in light curve properties and spectral evolution (see e.g. Krisciunas et al. 2006, A14), or to SN Ia colour and SED templates.

In Figure 5 the H07 template is plotted (dashed red) for the optical colours. Although the H07 template does extend from the UV to NIR, the data it is based on are sparse at the endpoints of this range. N08 studied the intrinsic optical colours between  $-10$  and  $+50$  days from  $B$ -band maximum. Two of their colour laws are shown (in blue) for normal,  $s_B = 1$ , SNe Ia in Figure 5 together with the colour dispersions they derive. Although these templates provide excellent coverage at optical wavelengths none of them cover, to high accuracy, the full wavelength range required for the UV-NIR analysis in this work.

The best studied unreddened SN Ia to date is SN 2011fe, discovered by the Palomar Transient Factory in the nearby spiral galaxy M101 (PTF11kly, Nugent et al. 2011). Its close proximity allowed detailed spectroscopic and photometric observations over a broad wavelength range from the UV (Brown et al. 2012; Mazzali et al. 2014), through the optical (e.g. Munari et al. 2013) to the near- (Matheson et al. 2012; Hsiao et al. 2013), and mid- (McClelland et al. 2013), infrared. The SN was also targeted in the far-IR (Johansson, Amanullah, & Goobar 2013) and radio (Chomiuk et al. 2012b), but was not detected at these wavelengths.

SN 2011fe is a normal SN Ia in both the optical and the NIR (Munari et al. 2013; Matheson et al. 2012) and does not show any spectroscopic peculiarities (Pereira et al. 2013). Further, the low Galactic and host galaxy reddening along the line-of-sight,  $E(B - V)_{\text{MW}} = 0.011 \pm 0.002$  mag and  $E(B - V)_{\text{host}} = 0.014 \pm 0.002$  mag (deduced from the integrated equivalent widths of the Na I D lines; Patat et al. 2013), makes it an excellent comparison object for studying reddening of SNe Ia. The colours of SN 2011fe shown in Figure 5 (solid black lines) have been obtained by combining the available UV-NIR data, as described in Appendix B based on the measurements (black triangles). We have also combined the spectroscopic data to create a daily sampled SED, shown in Figure B1. From the excellent  $B - V$  match between SN 2011fe and H07 we can also conclude that any potential discrepancies introduced by using the *SNooPy* model for the  $V$ -band is not likely to have any significant effect on the derived reddening laws when the colours are compared to the corresponding SN 2011fe colours.

In order to properly compare the colours of reddened SNe Ia to SN 2011fe, and fit reddening laws, we also need accurate estimates of the expected intrinsic colour dispersions. It is desirable to take into account how these vary in time, and how they are correlated both in time and with different colours. However, since we lack this information for all colours involved in the analysis, we adopt a simplified approach where all colours are treated equally with phase-independent dispersions for each colour. We further follow the procedure from A14 and assume that the colour uncertainties between SN phases are completely correlated and that there is no correlation between different colours.

Milne et al. (2013, M13 from hereon) following the work of Brown et al. (2010) and Milne et al. (2010), studied the UV–optical colours of 23 SNe Ia observed by *Swift*. They reported a dispersion of  $\sim 0.3$  mag in  $UV-v$  for their  $E(B - V) < 0.2$  mag sample after correcting the data for extinction and the result does not seem to be affected significantly by the choice of  $R_V$ . We adopt this dispersion, shown as grey bands around the SN 2011fe colour template in Figure 5, for the  $uvm2 - V$ , F225W  $- V$ , and F275W  $- V$  colours. For the F336W  $- V$  and  $U - V$  we adopt a dispersion of 0.1 mag based on the  $U - V$  dispersion from N08. Note that since these dispersions were derived for unreddened objects, and the dispersion of SNe Ia colours decrease with wavelength, they can be considered as conservative for highly reddened SNe. When reddened SNe Ia are studied using the *WFC3/UVIS* filters, the observations will effectively probe redder wavelengths compared to similar studies of unreddened objects as illustrated in e.g. Figure D1, and as a consequence the intrinsic dispersion of the former observations can be expected to be lower than the latter. On the other hand, even larger dispersions have been also observed in the UV, and we will use these intrinsic uncertainties under the assumption that all the SNe we study are similar to the reference SN 2011fe, which we will discuss further in §6.1.

B14 derive a dispersion for the pseudo-colour  $B_{\max} - V_{\max}$  of 0.06 mag using a Cauchy Prior for the colour distribution of their SNe Ia observed with the Carnegie Supernova Project. This is consistent with the phase-dependent results of N08 who derive a  $B - V$  dispersion of  $\lesssim 0.1$  mag, which we use here, within the range  $-10$  to  $+35$  days. For the  $V - R$  and  $V - i$  colours we adopt the N08 dispersions of 0.08 mag.

The NIR colours of SNe Ia typically show low dispersion (see e.g. Meikle 2000) and for the NIR-optical  $V - J$  and  $V - H$  colours we can estimate the dispersion in our phase range by using all SNe with  $E(B - V) < 0.2$  mag from the “Gold Sample” in Stanishev et al. (2015) and find 0.12 mag and 0.13 mag respectively. For  $V - K_s$  we conservatively adopt 0.2 mag based on the studies carried out at maximum (B14) and the dispersion of the  $K_s$ -band (Friedman 2012; Friedman et al. 2014).

The observed colour excesses,  $E(X - V)$ , for each reddened SN, together with the adopted colour dispersions, shown as background regions, are plotted in Figure 6. SN 2012bl has not been plotted in the figure. It will be omitted from the extinction analysis due to its lack of reddening.

If the colour excess is observed at single wavelengths, i.e.  $E(\lambda_1 - \lambda_2) = A_{\lambda_1} - A_{\lambda_2}$ , it will be time-independent for interstellar dust extinction. However, for broad-band pho-

tometry, we can expect  $E(X - V)$  to vary with time which is illustrated by e.g. the UV colours in Figure 6. As the intrinsic colour of the source varies with time, the effective wavelengths of all broad-band filters will change as well. Further, since the first term on the right-hand side in Equation (2) is also affected by a reddening law, any change in the effective wavelength of the filters  $X$  and  $V$  will affect this term more than the second term, which will induce a time-variability for  $E(X - V)$ . For most filters, the time-variability will be negligible except in the UV (where the extinction has a steep wavelength dependence), and in particular for filters with *red tails*, e.g. F225W (see Appendix D), the effect is significant. This is illustrated by the bottom panel of Figure D1 where the effective wavelength of three *WFC3/UVIS* UV filters have been plotted for a SN Ia at maximum for different extinctions.

## 4.2 Reddening laws

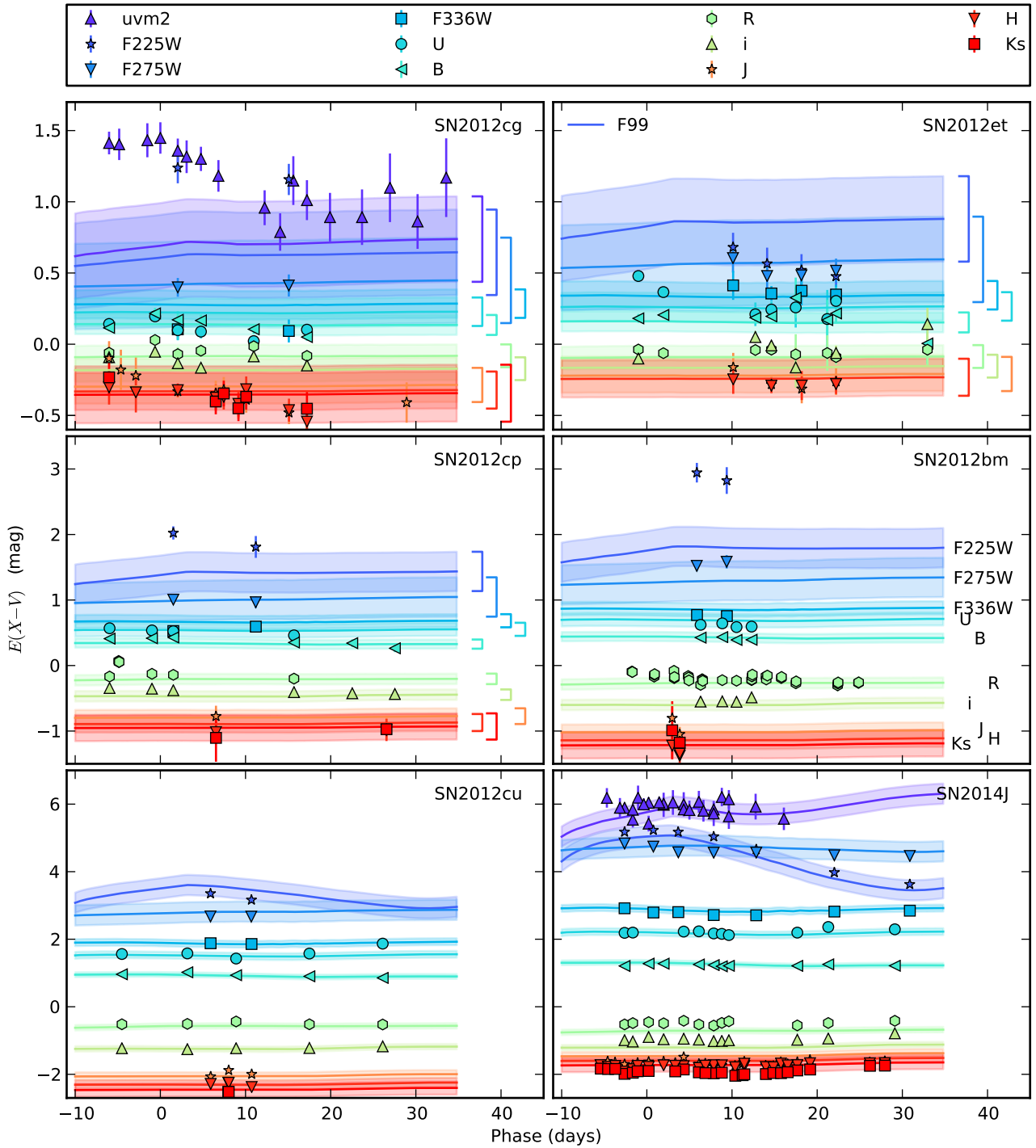
We test three different extinction law parametrisations. In addition to the widely used CCM law modified by O’Donnell (1994, hereafter CCM+O), which has been derived from studying different lines-of-sight in the Milky Way, we also fit the parametrisation from F99 that has been derived in a similar manner. For each of these we fit both the colour excess  $E(B - V)$ , which relates to the optical depth, and the ratio of the total-to-selective extinction,  $R_V$ . In the Milky Way the value of  $R_V$  typically varies between  $R_V = 2.2 - 5.8$ , for different lines-of-sight (e.g. CCM, Fitzpatrick 1999), despite this we allow extrapolations of  $R_V$  within the range  $R_V = 0.5 - 8$  when we fit the laws to the SN colours.

We also test a simple power-law model (Goobar 2008),

$$A(\lambda; A_V, a, \beta) = A_V \left[ 1 - a + a \left( \frac{\lambda}{\lambda_V} \right)^\beta \right],$$

where the reference wavelength,  $\lambda_V$ , will be chosen as  $\lambda_V = 0.55 \mu\text{m}$ . The CCM+O and F99 laws can be approximated by a power-law in the optical and NIR range, and Goobar (2008) showed that the observed reddening law of an object embedded in circumstellar dust can be approximated by the expression above with  $a = 0.8$  ( $a = 0.9$ ) and  $\beta = -2.5$  ( $\beta = -1.5$ ) for Milky Way (LMC) like dust. The reason why the original extinction laws are not preserved in the CS scenario is due to the geometry where multiply scattered photons on the CS dust will reach the observer while this is extremely unlikely for an ordinary interstellar dust geometry. Note that the values in Goobar (2008) were obtained under the assumption that the light source is constant in flux and colour. For time-dependent light sources the observed reddening is expected to vary in time which has been investigated for SNe Ia by e.g. Amanullah & Goobar (2011).

When the power-law is fitted to colour excesses,  $E(X - V) = A_X - A_V$ ,  $A_V$  and  $a$  are almost completely degenerate, and we are in fact only sensitive to the product  $a \cdot A_V$ . We will for this reason fix  $a$  to  $a \equiv 1$  in all fits below, which also has the implication that the resulting  $A_V$  can no longer be interpreted as solemnly being the extinction in the  $V$ -band. We have further checked that fixing  $a$  does not affect the fitted values of  $\beta$ , and that we obtain consistent values of the product  $a \cdot A_V$  when  $a$  is allowed as a free parameter,



**Figure 6.** The observed colour excess,  $E(X - V)$ , against phase for all SNe and filters. The predicted colour excess from the SN 2011fe template and the fitted F99 law, presented in Table 5, is shown with solid lines and the regions surrounding them mark the adopted intrinsic colour uncertainties. The uncertainties are also indicated by open brackets for the three SNe with the least reddening. The phase,  $p = (t - t_B^{\max})/s_B$ , was calculated based on the fitted  $t_B^{\max}$  and  $s_B$  values from Table 5. The apparent  $E(X - V)$  time-dependence of the F225W filter for e.g. SN 2014J originates from the red tails of these filters which causes the observed colour excess to vary with the intrinsic colour of the source.

although the individual values of  $a$  and  $A_V$  are in this case of course poorly constrained.

In A14, the SALT2 (Guy et al. 2007) law from Betoule et al. (2014) was also studied. This is a colour law and not an extinction law, and is further only defined for wavelengths  $\lambda < 8000 \text{ \AA}$ . It can therefore not be compared on an equal footing to the other laws and will not be studied in detail in this work. We will however compare it to the reddening or our low extinction SNe in § 6.5.

### 4.3 Results

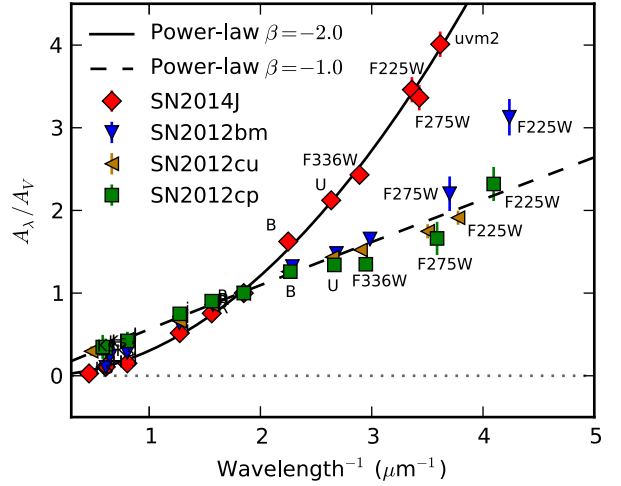
The fitted parameters for the three extinction laws are presented in Table 5 for each individual SN. The extinction laws were fitted to all available colours between phases  $-10$  and  $+35$  days and the predicted colour excesses using the best fitted F99 laws are also shown in Figure 6. The revised photometry of SN 2014J and updated intrinsic colour model with respect to A14 only had a minor impact on the fitted results for this highly reddened SN, and the values presented here are consistent with A14.

Further, all three reddening laws capture the general wavelength dependence of the observed colour excesses with the possible exception of SN 2012cg which will be discussed below. This shows that the observed SN colours from UV to NIR can indeed be described by the colours of SN 2011fe, with the adopted intrinsic dispersions, together with an extinction law with only two free parameters.

The six SNe span a broad range both in reddening,  $E(B - V) = 0.2\text{--}1.4$  mag, and in  $R_V = 1.4\text{--}3.0$  which confirms the findings from previous optical and NIR studies (e.g. B14), that point to a diversity of observed reddening laws of SNe Ia. The significant difference in the derived extinction laws is also illustrated in Figure 7 where the relative extinction,  $A_X/A_V$ , for the F99 law has been plotted for two different values of  $R_V$ . The diversity is particularly striking blue-wards of the  $U$ -band ( $\lambda^{-1} > 2.5 \mu\text{m}^{-1}$ ), emphasising the power of using UV data to study diversity in extinction.

In the Figure 7 we also show the relative extinction,  $A_X/A_V = 1 + E(X - V)/A_V$ , for the four SNe with  $E(B - V) > 0.2$  mag against inverse wavelength. Here,  $E(X - V)$  is calculated as the weighted average for all phases and  $A_V$  was obtained from the fitted F99 parameters as  $A_V = R_V \cdot E(B - V)$ . Because of this, the derived values for  $A_X/A_V$  are model-dependent and will be different for the different extinction laws. Further, since the extinction  $A_X$  is obtained from comparing observations in the filter  $X$  of reddened and unreddened sources, and the measurements correspond to different effective wavelengths, it is non-trivial to derive a specific wavelength corresponding to  $A_X/A_V$ . For these reasons, we follow the procedure from A14 and plot the  $A_X/A_V$  points in a manner where we chose to preserve the residuals between the fitted model and the data points from Figure 6. This is obtained by plotting each point,  $A_X/A_V$ , at the wavelength where the residual to the extinction law is the same as the corresponding weighted average residual of  $E(X - V)$  over all epochs. We emphasise that this plot is only a way of visualising the data and the reddening laws and that all extinction parameters were obtained by minimising Equation (3), i.e. not by fitting the extinction laws to the data points in Figure 7.

One observation that can be made by comparing the re-



**Figure 8.** Similar to Figure 7 but here the power-law fits from Table 5 are shown with the observed reddening. The wavelengths of the data points will differ from Figure 7 due to the way these plots are created, as explained in the text. Note that all fitted parameters described in this work were obtained by minimising Equation (3), and *not by fitting the extinction laws to the data points in this plot.*

sults in Table 5 and Figure 6 is that although all extinction laws appear to be general enough to describe all SNe, the F99 law shows tension for SNe 2012bm and 2014J. These SNe have in common that they both show a fair amount of reddening ( $E(B - V) > 0.5$  mag). The more dust along the line-of-sight to a SN, the more sensitive the fit will be to the extinction law, and less to possible intrinsic SN colour variations. In other words, we can expect any deviation between the assumed parametrisation and the actual observed extinction law to become more significant the redder the SNe are. It has already been pointed out that since the F99 law is empirical in nature, it may have features that do not properly extrapolate to low values of  $R_V$ . This could be the case for SN 2014J, where the F99 law with  $R_V = 1.4$  leads to a more prominent feature around  $\sim 1.5 \mu\text{m}^{-1}$  ( $\sim 6700 \text{ \AA}$ ) than what is supported by the data. In Figure 8 we compare the power-law fits for the four most reddened SNe, and here we can see how this is consistent with all the data points over the full wavelength range for SN 2014J.

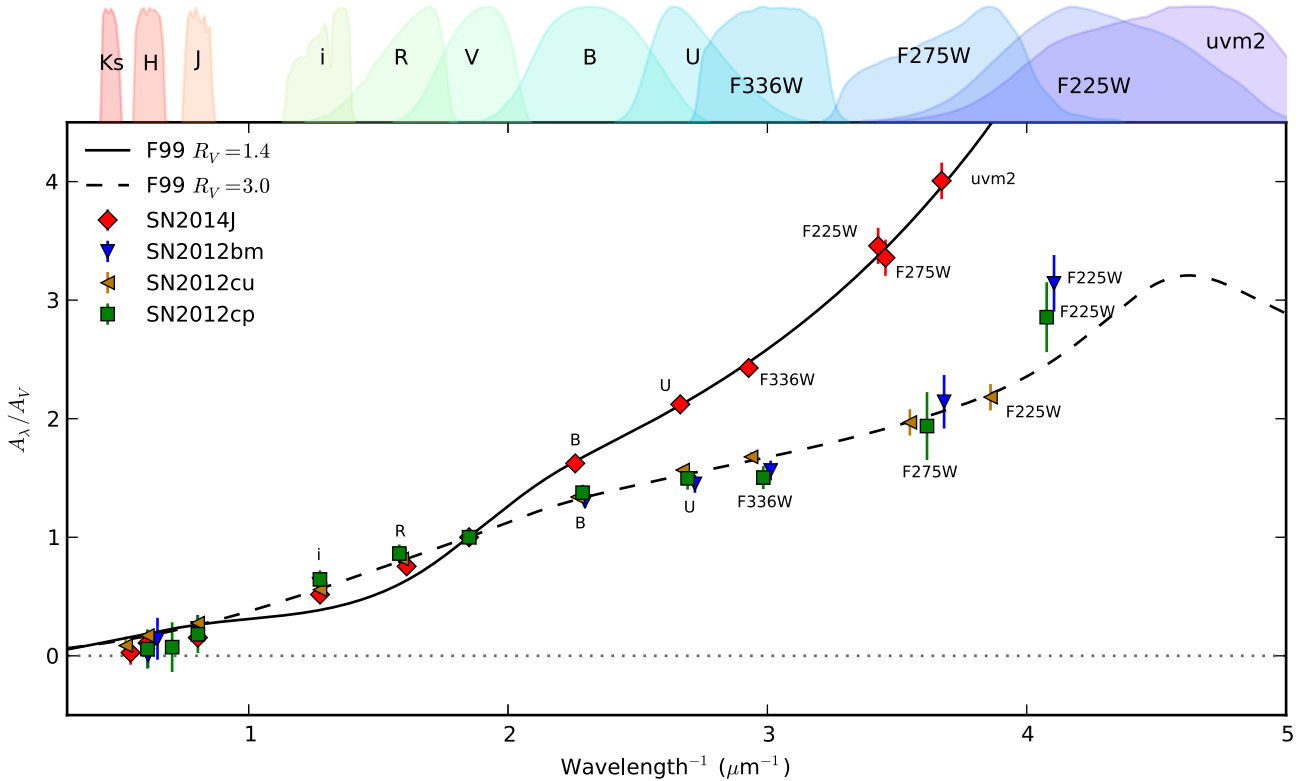
From the reduced  $\chi^2$  in Table 5 we can conclude that the power-law model provides better fits than F99 for both low- $R_V$  SNe, 2012et and 2014J, while the opposite is true for the  $R_V \lesssim 3$  SNe 2012cg, 2012cp and 2012cu. In particular, the F99 and CCM+O law seem to provide nearly perfect fits to the observations of SN 2012cu, which is also seen in Figure 7. SN 2012bm shows tension for all fitted laws.

The CCM+O law performs better in terms of goodness-of-fit for several cases, e.g. for SN 2014J. However, this law does not extrapolate well to low  $R_V$  values which give rise to a sharp, discontinuous, knee when the NIR and optical parametrisations of the law are merged.



	2012cg	2012et	2012cp	2012bm	2012cu	2014J	
F99	$E(B - V)$	0.15 (0.02)	0.17 (0.03)	0.35 (0.03)	0.46 (0.04)	0.99 (0.03)	1.36 (0.02)
	$R_V$	$2.7^{+0.9}_{-0.7}$	$1.7^{+0.6}_{-0.5}$	$3.0^{+0.4}_{-0.4}$	$3.0^{+0.5}_{-0.4}$	$2.8^{+0.1}_{-0.1}$	$1.4^{+0.1}_{-0.1}$
	$t_B^{\max}$	56080.6 (0.3)	56191.0 (0.5)	56081.1 (0.3)	56018.0 (1.9)	56104.5 (0.4)	56687.9 (0.2)
	$s_B$	1.12 (0.02)	1.01 (0.03)	1.18 (0.02)	1.24 (0.12)	1.04 (0.03)	1.14 (0.01)
	$\chi^2/\nu$	1.67	1.50	1.60	2.83	0.63	2.38
CCM+O	$E(B - V)$	0.15 (0.02)	0.17 (0.03)	0.36 (0.03)	0.48 (0.04)	1.00 (0.03)	1.40 (0.02)
	$R_V$	$2.6^{+0.8}_{-0.7}$	$1.4^{+0.8}_{-0.7}$	$3.0^{+0.4}_{-0.4}$	$2.8^{+0.4}_{-0.4}$	$2.8^{+0.2}_{-0.2}$	$1.2^{+0.1}_{-0.1}$
	$t_B^{\max}$	56080.6 (0.3)	56191.1 (0.5)	56081.1 (0.3)	56019.1 (1.7)	56104.7 (0.4)	56688.1 (0.2)
	$s_B$	1.12 (0.02)	1.01 (0.03)	1.18 (0.02)	1.18 (0.10)	1.03 (0.03)	1.13 (0.01)
	$\chi^2/\nu$	1.60	1.34	1.42	2.14	0.91	1.42
Power-law	$A_V$	0.44 (0.19)	0.30 (0.13)	1.51 (0.40)	1.47 (0.43)	3.59 (0.29)	1.97 (0.10)
	$\beta$	-1.2 (0.4)	-1.7 (0.5)	-0.8 (0.2)	-1.2 (0.4)	-1.0 (0.1)	-2.0 (0.1)
	$t_B^{\max}$	56080.6 (0.3)	56191.1 (0.5)	56081.2 (0.3)	56020.4 (2.7)	56104.8 (0.4)	56688.1 (0.2)
	$s_B$	1.12 (0.02)	1.00 (0.03)	1.18 (0.02)	1.16 (0.15)	1.02 (0.03)	1.13 (0.01)
	$\chi^2/\nu$	1.79	1.17	1.95	2.13	1.78	1.65

**Table 5.** The best fitted parameters to all measured colours between phases  $-10$  and  $+35$  days for the reddening laws investigated in this work. For the power-law we have fixed  $a$  to  $a \equiv 1$  to break the degeneracy between this parameter and  $A_V$ . For each value we quote the 68% level fitted uncertainty when the parameters are considered individually. The degrees of freedoms were calculated under assumptions that the colours are independent measurements, which is not the case.



**Figure 7.** The normalised extinction,  $A_X/A_V$  for the four SNe with  $E(B - V) > 0.2$  mag. Here,  $A_X/A_V = 1 + E(X - V)/A_V$  where  $E(X - V)$  is calculated as the weighted (with the measurement uncertainties) averages for all epochs, and  $A_V$  is obtained from the fitted F99 parameters in Table 5. Following A14, each data point has been plotted at the wavelength where the residual to the extinction law is matching the corresponding residual between  $E(X - V)$  and the fitted reddening law in Figure 6. Note that all fitted parameters described in this work were obtained by minimising Equation (3), and *not by fitting the extinction laws to the data points in this plot*. The filter passbands for the filters used for the minimisation are shown at the top.

SN	Phase	$E(B-V)$	$R_V$	$\chi^2/\nu$
SN2012cg	[-5, +5]	0.16 (0.02)	$1.7^{+0.8}_{-0.6}$	1.51
	[+10, +20]	0.13 (0.03)	$3.6^{+1.5}_{-1.0}$	1.34
SN2012cp	[-5, +5]	0.33 (0.03)	$1.8^{+0.6}_{-0.5}$	1.63
	[+10, +20]	0.34 (0.03)	$2.2^{+0.8}_{-0.6}$	0.55
SN2014J	[-5, +5]	1.35 (0.03)	$1.5^{+0.1}_{-0.1}$	1.81
	[+10, +20]	1.32 (0.03)	$1.6^{+0.1}_{-0.1}$	1.56

**Table 6.** Fitted F99 parameters using only data between  $-5$ – $+5$  and  $+10$ – $+20$  days from maximum, respectively. While SNe 2012cg and 2014J both have UV–NIR observations for both intervals, the fits to SN 2012cp are only based on UV–optical data. The parameters  $t_B^{\max}$  and  $s_B$  were fixed to the values given in Table 5.

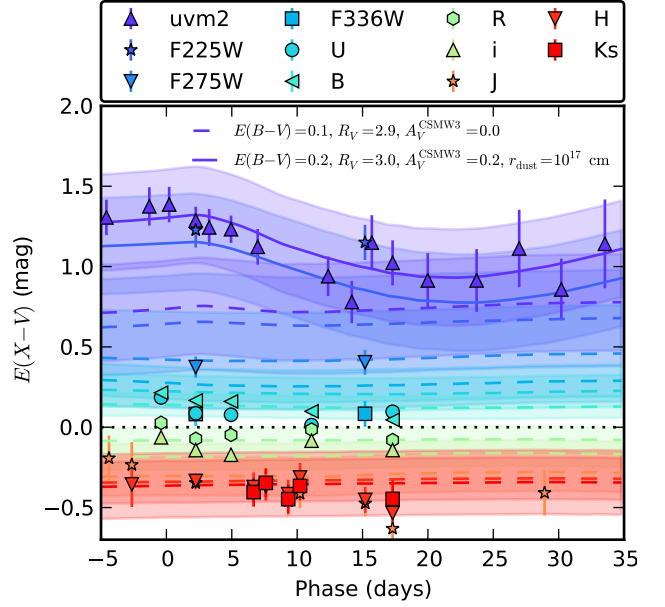
## 5 SEARCHING FOR CIRCUMSTELLAR DUST

All the SNe studied here show a colour excess evolution that is within the adopted dispersion with the exception of SN 2012cg. This SN reveals a minimum in its NIR colour evolution around two weeks past max which is possibly accounted to the diversity observed in NIR SN Ia light curves between the first and second peaks. This is within the assumed intrinsic uncertainty although the assumption that these are fully correlated between epochs does not hold in this case. What is perhaps even more striking are the  $\sim 2\sigma$  variations of the  $uvm2 - V$  evolution which suggests that this colour evolve slower than the corresponding 2011fe colour, resulting in an apparent decrease with time of the colour excesses.

The fact that the observed colour evolution of SN 2012cg is not consistent with a single extinction law is also illustrated by Table 6 where the F99 law has been fitted for two phase intervals: around maximum and between 10 and 20 days past  $t_B^{\max}$ . The fits were carried out for the three SNe that had UV data in both these intervals while keeping the fitted values of  $t_B^{\max}$  and  $s_B$  fixed to the values given in Table 5. The fitted parameters of SNe 2012cp and 2014J are within errors between the two epochs while this is not the case for SN 2012cg.

Evolving colour excesses, or a reddening law that changes with time, is predicted if circumstellar dust is present in the SN environment. Amanullah & Goobar (2011) studied this effect at optical wavelengths and we can extend that analysis, using the same tools, to also cover UV and NIR in order to test the observed colour evolution of SN 2012cg against the expectations for a CS scenario.

In Figure 9 the same colour excesses of SN 2012cg that are plotted in the upper left panel of Figure 6 are shown. The solid lines represent the result when we fit a combined F99 and CS dust law to the bluest colours,  $uvm2 - V$  and  $F225W - V$ , while keeping  $t_B^{\max}$  and  $s_B$  fixed to the values from Table 5. The best fit values for the F99 law are  $E(B-V) \approx 0.2$  mag and  $R_V \approx 3.0$ . The CS dust in this case consists of dust modeled to match the average extinction properties in the Milky Way (Draine 2003) and distributed



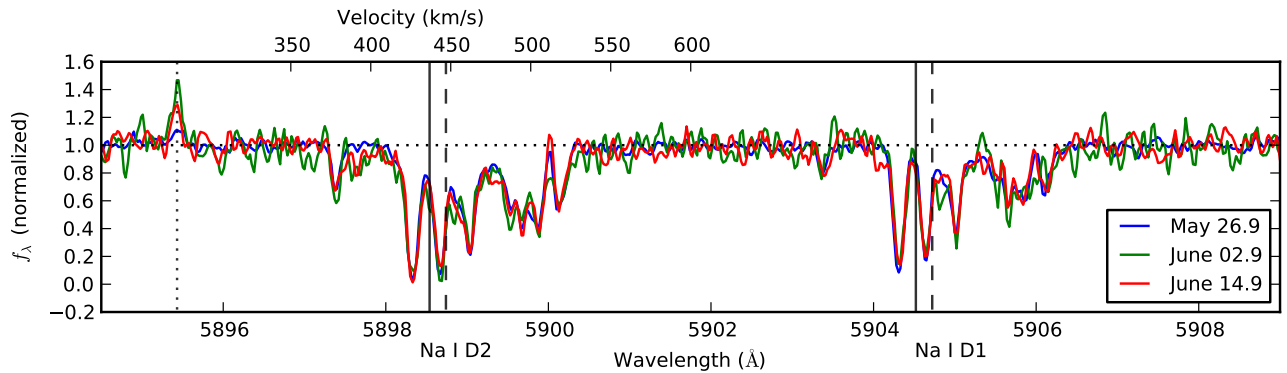
**Figure 9.** The measured colour excesses for different phases for SN 2012cg are taken from Figure 6. Overplotted are the best fits of a combined circumstellar and F99 dust law, where the former has been adopted from Amanullah & Goobar (2011) where the dust is distributed in a thin spherical shell. The solid lines show the best fit when only the  $uvm2 - V$  and  $F225W - V$  colours are considered, while the dashed lines show the fit to all colours. The former is obtained with a CS dust shell at radius  $r_{\text{dust}} \sim 10^{17}$  cm from the SN that give rise to an extinction of  $A_V \sim 0.2$ , while the latter fit is consistent with no CS dust.

in a thin spherical shell. Here the best fit yields an extinction of  $A_V^{\text{CSMW3}} \approx 0.2$  for a shell at a radius of  $r_{\text{dust}} \sim 10^{17}$  cm.

Although adding Milky Way-like CS dust to the picture can give an explanation for the time-evolution of the bluest colours, it does not provide an improvement beyond a single F99 law when all colours are considered simultaneously, which is shown by the dashed lines in Figure 9. Here the best fit is obtained for  $A_V^{\text{CSMW3}} \approx 0.0$ . The explanation for this is that the specific CS dust model we test also predicts evolution of the redder colours which in this case is not supported by the data. If CS dust is present, the total extinction,  $A_\lambda$ , can only decrease with time due to late-arriving scattered photons. From the results in Table 6, we see that this is not the case for  $A_V$ , although the value of  $A_V$  is in this case model dependent.

It is also possible to detect CS dust by studying the time-evolution of multi-epoch high-resolution spectroscopy. This was mentioned in § 1 and is the observational signature that has been used to claim detections of CS dust around SNe Ia.

For SN 2012cg we obtained high-resolution FIES spectra for three epochs which are listed in Table 4 and correspond to phases  $-7$ ,  $1$  and  $13$  days with respect to  $t_B^{\max}$ . All spectra contain well resolved unsaturated Na I D features (Figure 4), which are shown in Figure 10. The bulk of the distinguishable features are likely part of the visible interstellar medium (ISM), but it can not be excluded that Na I associated with CS medium could contribute to the observed



**Figure 10.** Na I D absorption doublet of SN 2012cg. The rest frame wavelength of Na I D in NGC 4424 (solid lines) and for the approximate stellar velocity along the line-of-sight (dashed) are indicated. The dotted vertical line indicates a sky emission line. In Section 5 the profile of Na I D is discussed in detail.

profile. A full description of the Na I D profile and summary of other interstellar absorption features is presented § 6.4.

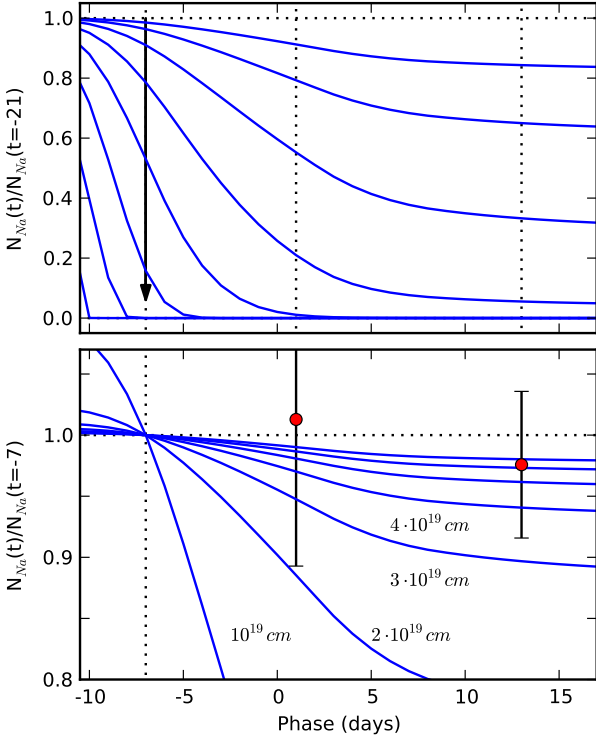
The FIES spectra were normalised by fitting 3rd-order polynomials to the continua bracketing Na I D. No telluric corrections were performed on the spectra, due to the lack of standard star observations on some of the epochs. We fitted Gaussian profiles to the weighted average of the three epochs, and find that the Na I D line ratios (D1/D2) range from 2–1.2 for individual features. The deepest features have low ratios, indicating that they are not optically thin. Based on a standard star spectrum taken on the second epoch, we identify telluric features overlapping with Na I D2, where as Na I D1 appears to be located at a less affected part of the spectrum. We therefore focus most of the time-variation analysis on Na I D1, because it is less contaminated by telluric features and it is more likely to be optically thin due to the lower absorption cross-section of this transition. We measure the total equivalent width of Na I D1 of the respective epochs to be,  $702 \pm 9$ ,  $711 \pm 32$  and  $685 \pm 20$  mÅ.

Since no significant time-variations of Na I D absorption between different epochs could be detected beyond the noise level, we attempt to use the non-detections to constrain possible CS dust models assuming a thin sphere of CS material shell distribution around the SN. Using the Na I photo-ionisation model described by Borkowski, Blondin & Reynolds (2009) and recently applied to SN 2014J in Graham et al. (2015b), we can exclude CS material, and thus also CS dust, at certain radii from the SN. If the Na I is optically thin, a change in equivalent width of Na I D is directly proportional to a decrease in column density. Furthermore, the product of the integrated photon count flux above the ionisation energy of Na I (2400 Å) and its corresponding ionisation cross-section is proportional to the fractional decrease in column density. We can thus compute the fractional decrease in equivalent width of Na I D expected due to photo-ionisation of CS material at a given radius from a SN. In this model, CS Na I can be excluded which is far enough away from the SN to not have been fully ionised by the first epoch and close enough to see changes by the following epochs. The diversity of SNe Ia spectra in UV implies that the results will be sensitive to the spectral templates used. Noting that the *STIS* spectra of SN 2012cg resemble those

of SN 2011fe, we use the SN 2011fe spectral template described in Appendix B. However, the H07 template extends to shorter wavelengths and earlier times than the SN 2011fe template. We therefore extrapolate our spectral template by H07, whereby the flux is rescaled for continuity. This slightly increases the photo-ionising flux.

In Figure 11 the model fractional decrease of Na I at different radii from a SN are shown. We determined the inner radius of excluded CS Na I of 2012cg by assuming that an absorption feature deeper than three times the root-mean-square of the noise would be detectable in the earliest spectrum. This yields a distance of  $r_{\text{dust}} < 2 \cdot 10^{18}$  cm within which the Na I column density would have decreased below detection levels since explosion. With respect to the first epoch, the fractional change of the equivalent width of Na I D1 is considered to set an upper limit on the change in column density. Within  $3\sigma$  errors, the last epoch excludes the models in which the dust is located closer than  $r_{\text{dust}} > 3 \cdot 10^{19}$  cm from SN 2012cg. Using the total equivalent width of Na I D1 of course has the caveat that it considers the entire measured column density to be situated at one radius from the SN, which we know cannot be true based on the profile. We therefore also consider how much a single feature must change to be detected given the signal-to-noise of the spectra. The limits obtained by the absence of change beyond three times the root-mean-square noise are comparable to those set by the fractional change of the total equivalent width. Lastly we note that the limits are not valid if there is CS material that is not optically thin, since the equivalent width then does not change linearly with column density.

In summary, although the existence of Milky Way-like CS dust at  $r_{\text{dust}} \sim 10^{17}$  cm is consistent with the observed time-evolution of  $E(uvm2 - V)$  and non-detection of time-varying Na I D absorption for SN 2012cg, the minimal  $E(X - V)$  evolution of the redder wavelengths disfavors this explanation. This suggests that CS dust can only be present either in very small amounts,  $A_V^{\text{CSMW3}} < 0.05$  or if it is located at larger radii  $r_{\text{dust}} \sim 10^{19}$  cm for which the time-delays are much larger than the time-scales studied here and are not expected to affect the observed colours (Amanullah & Goobar 2011). Further, non-detections in ra-



**Figure 11.** Na I ionisation models for thin shells of CS material at different distances from a SN, suggesting the absence of circumstellar material at distances between  $r_{\text{dust}} = 2 \cdot 10^{18} - 3 \cdot 10^{19}$  cm around SN 2012cg. The upper panel shows the fractional decrease of Na I due to ionization as a function of time at radii ranging from  $r_{\text{dust}} = 10^{18} - 3 \cdot 10^{19}$  cm. The arrow indicates an approximate upper limit for Na I around 2012cg to be ionised beyond detection level by  $-8$  days from  $B$  maximum based on the S/N around the Na I D doublet of the first FIES spectrum. The lower panel shows the fractional change of Na I with respect to the first epoch. The fractional change of equivalent width of Na I D1 with  $3\sigma$  error bars set an approximate lower radius limit at which CS material could be present while the ionisation effects remain undetected.

dio (Chomiuk et al. 2012a) and far-IR (Johansson, Amanullah & Goobar 2013) are other observational signatures suggesting that SN 2012cg exploded in an environment free from CS material. Here the far-IR observations are sensitive to pre-existing CS dust that would be heated by the explosion.

## 6 DISCUSSION

### 6.1 Using SN 2011fe as a colour template

One possible explanation for the time-evolution of the SN 2012cg  $UV - V$  colour excess could be that it is intrinsically different from SN 2011fe. We have already argued that all the SNe in our sample, are normal SNe Ia by comparing their overall spectra (see Figure 4) and their SNID classification, but we can also study individual features which has proven to be a useful approach for sub-classifying normal SNe Ia.

The properties of individual features can be quanti-

SN	Phase (days)	$pW1$ (Å)	$pW6$ (Å)	$pW7$ (Å)	Type
2011fe	+0	94 (1)	14 (1)	94 (1)	CN
2012bl	-1	110 (2)	1 (1)	89 (2)	CN
2012bm	+9	125 (2)	13 (1)	76 (2)	CN
2012cg	+2	103 (4)	9 (3)	75 (4)	CN
2012cp	+2	98 (1)	10 (1)	74 (1)	CN
2012cu	+9	70 (1)	24 (1)	99 (1)	CN
2012et	+2	123 (2)	11 (1)	133 (2)	BL
2014J	-1	155 (2)	14 (1)	100 (3)	CN

**Table 7.** Calculated pseudo-equivalent widths ( $pW$ ) based on the optical spectra. The  $pW$  naming convention from Garavini et al. (2007) is followed with  $pW1$  corresponding to Ca II H&K, and  $pW6$  and  $pW7$  corresponding to Si II  $\lambda 5972$  Å and Si II  $\lambda 6355$  Å respectively. The quoted uncertainties are the statistical errors. Systematic uncertainties from estimating the pseudo continuum could be dominating in some cases. The classification as "core normal" (CN) or "broad-line" (BL) is based on the criteria from Blondin et al. (2012). Given that these criteria were derived based on spectra around maximum, the classification of SN 2012cu should be considered less certain. The measurement of SN 2014J was carried out on a spectrum published in Goobar et al. (2014). See the text for details.

fied by e.g. measuring their velocities and pseudo-equivalent widths ( $pW$ ), where pseudo refers to the fact that the equivalent widths are obtained using a pseudo continuum since proper continua are absent in SN Ia spectra. The pseudo continuum can be defined with a straight line between the two flux peaks surrounding an absorption feature and the  $pW$  is then calculated as the integral of the spectrum flux relative to the continuum. Using this method, the error of the measurement will typically be dominated by the systematic uncertainty introduced by the determination of the pseudo-continuum.

Branch et al. (2006) suggested a classification scheme based on pseudo-equivalent width measurements of the absorption features near  $5750$  Å and  $6100$  Å (following the naming convention from Garavini et al. (2007) we refer to these as  $pW6$  and  $pW7$  respectively), that can be associated with Si II  $\lambda 5972$  Å and Si II  $\lambda 6355$  Å as shown in Figure 4. They identify four different groups when considering  $pW6$  and  $pW7$  for their sample where in particular the "core normal" (CN) SNe, to which SN 2011fe belongs, are tightly clustered and show a high degree of general spectral homogeneity.

In Table 7 we show measurements of  $pW6$  and  $pW7$  for spectra close to maximum after correcting our spectra for host galaxy reddening (although this is not expected to have any major impact on the results as discussed in Nordin et al. 2011) using the fitted reddening laws from §4.3. We also present the  $pW$  for Ca II H&K ( $pW1$  following Garavini et al. 2007) and note that all SNe have  $pW1 < 150$  Å. It has been shown that  $pW1$  correlates with the intrinsic  $U - B$  colour (Folatelli et al. 2013; Burns et al. 2014) where SNe with  $pW1 > 150$  Å are intrinsically redder than the bulk of normal SN Ia.

Using the classification criteria from Blondin et al. (2012) we can classify all SNe as CN except for SN 2012et which lands in range that Branch et al. (2006) defines as the "broad-line" (BL) group. These SNe Ia have broader and

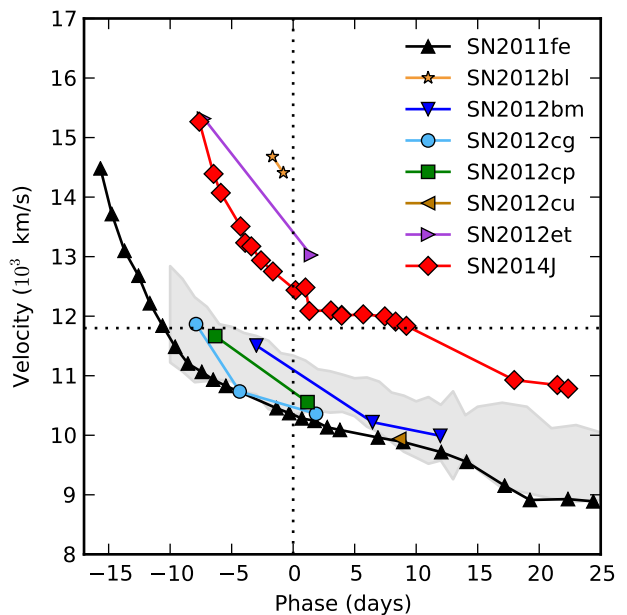
deeper 6100 Å absorption but are in most aspects not very different from CN SNe. When the observed spectra from the two groups are compared with synthetic spectra generated with SYNOW (Branch et al. 2003) similar photospheric velocities and excitation temperatures can be used to describe both the CN and BL groups (Branch et al. 2006).

Given these results, there is no evidence for any significant discrepancy between the optical properties of our SNe and SN 2011fe, with the exception for the additional source of luminosity observed in the very early optical lightcurve of SN 2014J (Goobar et al. 2015). However, several studies have shown that normal SNe in the optical still show dispersion in the UV and M13 argued that their sample could be divided in up to four groups. Here, SN 2011fe belongs to what they define as the "near-UV (NUV) blue" group which they conclude is on average 0.44 mag bluer in  $u - v$  than the reddest, "NUV-red", group. They consider SNe with  $E(B - V) < 0.25$  mag, and this constraint could still allow extinction to contribute with up to 0.9–1.3 mag (depending on the extinction law) to the  $uvm2 - v$  colour, which may offer an explanation for the significant colour dispersion. However, M13 show that the two groups are distinguishable even after the individual SNe have been corrected for reddening.

Further, they also conclude that the groups, on average, show different spectroscopic properties. All NUV-blue SNe have "normal" photospheric expansion velocities around maximum when quantified based on the Si II  $\lambda 6355$  Å feature, while the remaining groups are populated with both "normal" (NV) and "high" (HV) velocity SNe. They use the definition from Wang et al. (2009), where SNe with Si II velocities of  $v_{\text{Si II}} > 11,800 \text{ km s}^{-1}$  at maximum are classified as HV and the remaining as NV. Although there may not be a sharp distinction between the two classes (see e.g. Silverman, Kong & Filippenko 2012), the M13 results are consistent with previous studies showing that an optical colour-velocity dependence exists with HV SNe being on average intrinsically redder (Foley & Kasen 2011; Blondin et al. 2012; Folatelli et al. 2013; Mandel, Foley & Kirshner 2014). On the other hand, exceptions from these findings is, SN 2011de, the brightest SN Ia observed (Brown 2014), and the results by Wang et al. (2012), showing, using the *HST* Advanced Camera for Surveys, that the UV-optical colours of the HV object SN 2004dt are significantly bluer than the NV SN 2005cf.

We have measured the Si II velocities, which are presented in Figure 12, for all spectra we have obtained of the SNe at  $t < +25$  days from  $t_B^{\text{max}}$  by fitting Gaussian profiles. From the figure we conclude that SNe 2011fe, 2012bm, 2012cg and 2012cp show similar velocity evolutions, and the single measurement we have of SN 2012cu is also consistent with SN 2011fe. The SNe 2012bl, 2012et and 2014J, on the other hand, are classified as HV. Note however that the velocity range for the full sample is still small enough for the scatter induced by the optical colour-velocity relation to land within the adopted uncertainty of 0.1 mag for  $B - V$  (Mandel, Foley & Kirshner 2014). Mandel, Foley & Kirshner (2014) further find no evidence for such a relation for the  $V - R$  and  $V - i$ .

SN 2012et is both classified as BL (although the  $pW6$  and  $pW7$  values put its close to the border of the BL and CN groups) and HV, and such objects have been shown to



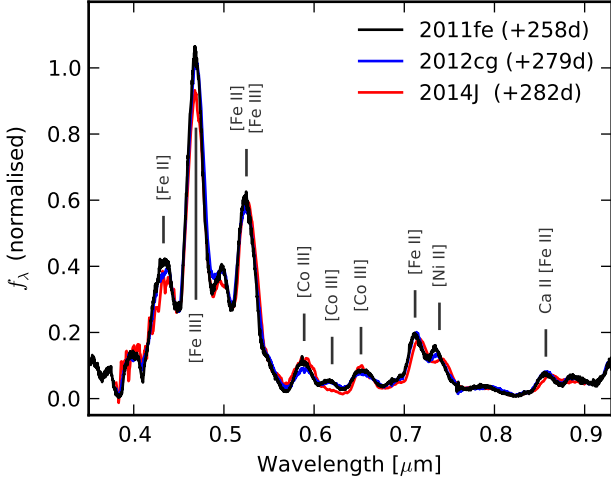
**Figure 12.** Measured velocities of Si II  $\lambda 6355$  for different phases of the SNe studied in this work together with reference objects. For SNe 2011fe and 2014J the velocities were calculated based on the spectra published by Pereira et al. (2013), Goobar et al. (2014) and Marion et al. (2015). The grey region is marking the average SN Ia velocity range (Folatelli et al. 2012, 2013). The dashed line marks the difference used to separate high and low velocity objects by Wang et al. (2009) at maximum.

be redder than the bulk of normal SNe Ia (Folatelli et al. 2013; Burns et al. 2014). However, the value of  $pW1$  for SN 2012et at  $-8$  days is  $pW1 = 82 \pm 7 \text{ \AA}$ , which is well below  $pW1 \sim 150 \text{ \AA}$  where the BL HV start to deviate in colour (see Figure 17 in Burns et al. 2014).

M13 also determine that NUV-blue SNe show evidence of unburned C II in their optical spectra while the remaining groups consist of objects both with and without unburned C II detections. These results are also consistent with optical spectroscopic studies where SNe Ia with signatures of unburned carbon are bluer in their optical colours (e.g. Thomas et al. 2011; Folatelli et al. 2012; Silverman & Filippenko 2012). Unburned carbon can be detected through absorption attributed to the most prominent line in the optical, C II  $\lambda 6580$ , but the prospects of detecting this line decreases for epochs approaching maximum light (see e.g. Figure 11 in Folatelli et al. 2012). The SNe 2011fe, 2012cg and 2014J have all evidence for unburned material in their spectra (Nugent et al. 2011; Silverman et al. 2012; Marion et al. 2015), while the first spectra of the remaining objects were all obtained  $> -7$  days with respect to  $t_B^{\text{max}}$ .

Finally, we can also compare the optical spectra of SNe 2011fe, 2012cg and 2014J at the late-time nebular phase when the material in the SN ejecta is optically thin. In Figure 13 we present spectra of the three SNe at 260–280 days past maximum after they have been corrected for reddening. The spectra are dominated by forbidden lines from the decay chain  $^{56}\text{Ni} \rightarrow ^{56}\text{Co} \rightarrow ^{56}\text{Fe}$ . The spectra of SNe 2012cg and 2011fe are remarkably similar while for SN 2014J the [Co III] line at 6200 Å is suppressed and the features at





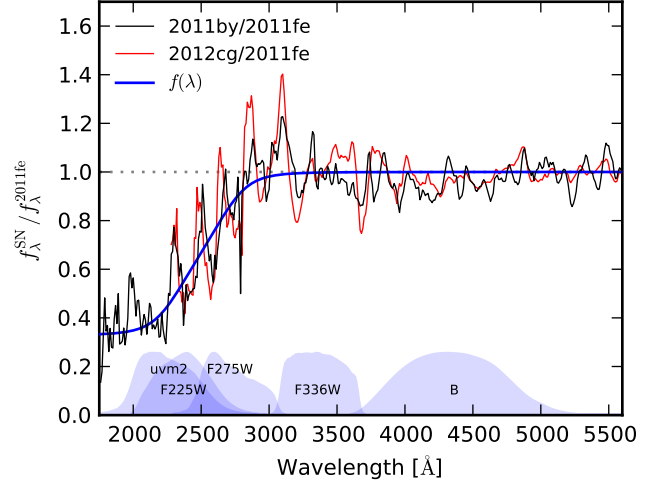
**Figure 13.** Nebular spectra of SNe 2011fe, 2012cg and 2014J obtained with the William Herschel Telescope, the Keck and the Apache Point Observatory. The epochs of the spectra with respect to maximum is specified in the legend. The spectra of SNe 2012cg and 2014J were corrected for reddening and smoothed using a 5 pixel boxcar, and then scaled to match the spectrum of SN 2011fe. The main features of the nebular spectra have been marked.

7000–7400 Å are redshifted with respect to the others. The two peaks identified as [Fe II]  $\lambda$ 7155 and [Ni II]  $\lambda$ 7378 (e.g. Maeda et al. 2010a,b), are consistently blueshifted with respect to their rest frame by 800–1000 km/s in SNe 2011fe and 2012cg. In contrast, these lines appear redshifted in SN 2014J by a similar amount. This is in line with the findings of Maeda et al. (2010a) that supernovae showing a HV gradient (Benetti et al. 2005), generally associated with HV supernovae (Wang et al. 2009), demonstrate exclusively redshifted nebular velocities, while LV supernovae preferentially demonstrate blue shifted nebular profiles.

In conclusion, based on the optical spectroscopy available both around maximum light and in the nebular phase, there are no observational signatures that would disfavour comparing the colours of the SNe to SN 2011fe for studying reddening together with the adopted uncertainties. In particular, while SNe 2011fe and 2012cg have slightly different light curve shapes, they show remarkable similarities in spectroscopical features and evolution. However, the SNe 2011by and 2011fe have also been found to be almost identical in the optical while showing significant discrepancies in the UV (Foley & Kirshner 2013; Graham et al. 2015a). Next we investigate the impact on the fitted reddening laws if we instead compare the measured colours to SN 2011by.

## 6.2 SN 2011by as colour template

SN 2011by was discovered (Jin & Gao 2011) in NGC 3972 and found to be a spectroscopically normal SN Ia (Silverman, Ganeshalingam & Filippenko 2013) with minimal reddening (Maguire et al. 2012). *HST/STIS* UV observations were obtained at maximum (Program GO-12298; PI: Ellis) and using these data Foley & Kirshner (2013) and Graham et al. (2015a) have shown that while SN 2011by is almost identical to SN 2011fe in its optical properties it is intrin-



**Figure 14.** The flux ratio between SNe 2011by and 2011fe, and 2012cg and 2011fe at maximum after they have been corrected for Galactic extinction, corrected for reddening (in the case of SN 2012cg), and normalised between 5000–5500 Å. We use the reddening parameters derived for maximum for SN 2012cg from Table 6 and both spectra have been smoothed using boxcar smoothing. These spectra are only shown for wavelengths where the signal-to-noise per spectral element exceeds 2.5. Some of the passbands have been plotted in the background to illustrate which filters are affected by the difference in spectroscopic shape. The blue line shows Equation (C1) that is used to model the flux ratio between SNe 2011by and 2011fe as described in Appendix C.

sically fainter in the UV. The two SNe are compared at maximum in Figure 14, where the ratio between the spectra are plotted. Comparing this to the ratio between SNe 2012cg and 2011fe at maximum, also shown in the figure, suggests that SN 2011by may be an even better reference SN than SN 2011fe in the UV.

Both SNe 2011by and 2012cg are fainter than SN 2011fe in the UV, which could explain why SN 2012cg is marginally redder in its UV colours than predicted by the fitted reddening laws. Two other SNe in the sample, 2012bm and 2012cp, show a similar behaviour for the bluest colour F225W – V.

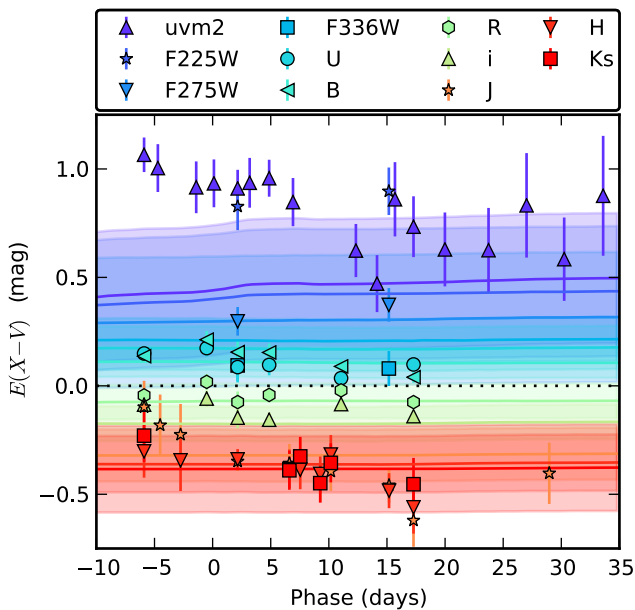
*HST/STIS* spectroscopy of SN 2011by was obtained at phases –10 and –1 days, which unfortunately does not allow us to construct a SED covering the full range of our observations as we did for SN 2011fe. However, we can use the available spectrum at maximum and the *Swift* observations of this SN to model the difference between SNe 2011by and 2011fe in the UV. In Appendix C we describe this procedure and construct a model SED of SN 2011by under the assumption that it is identical to SN 2011fe at wavelengths  $\lambda > 3000$  Å.

Using this as the comparison SN, we can refit the reddening laws for the SNe 2012cg, 2012cp and 2012bm for which we have reason to believe that SN 2011by may provide a better comparison object. The fitted F99 parameters at maximum are presented in Table 8. The fitted parameters for SNe 2012cp and 2012bm change marginally compared to Table 5 but we do note that reduced  $\chi^2$  has decreased. However, the most interesting result here is for SN 2012cg. Not only the fitted reddening law parameters change, but also



	2012cg	2012cp	2012bm
$E(B - V)$	0.12 (0.02)	0.34 (0.03)	0.45 (0.04)
$R_V$	$3.7^{+1.4}_{-1.0}$	$3.2^{+0.5}_{-0.4}$	$3.0^{+0.4}_{-0.4}$
$t_B^{\max}$	56080.5 (0.3)	56081.1 (0.3)	56018.3 (1.1)
$s_B$	1.13 (0.02)	1.19 (0.02)	1.24 (0.08)
$\chi^2/\nu$	0.88	1.09	2.13

**Table 8.** Best fitted parameters of the F99 law using the colour and SED model that has been adapted to match SN2011by in the UV as described in §C. The quoted uncertainties are the statistical  $1\sigma$  errors from the  $\chi^2$  fit.



**Figure 15.** The colour excess for SN 2012cg, similar to the upper left panel of Figure 6, with the exception that the measured and predicted colour excesses are based on the SED model described in §C, which has been derived to mimic SN 2011by in the UV. The solid lines represent the best fitted F99 law with the parameters given in Table 8.

the time-evolution of  $E(uvm2 - V)$  becomes less significant, which can be seen in Figure 15. This can be explained from Figure C1, where the colour excesses of SN 2011by are shown with respect to SN 2011fe, and a similar time-evolution for  $E(uvm2 - V)$  is shown. We note that there is still a tension between the bluest colours and the fitted reddening law which may suggest that the spectrum of SN 2012cg drops faster than SN 2011by at  $\lambda < 2500 \text{ \AA}$ .

From the spectroscopic similarities between SNe 2011by and 2012cg as shown in Figure 14 and the similar UV-optical colours we find that SN 2011by appears to be a better comparison SN for SN 2012cg than SN 2011fe. This study also shows how studying the UV-colour evolution of SNe Ia provides a powerful tool for comparing predictions of CS dust scenarios with intrinsic colour variations.

SN	$\lambda$ -range	$E(B - V)$	$R_V$	$\chi^2/\nu$
SN2012cg <sup>†</sup>	UV-NIR	0.11 (0.02)	$3.8^{+1.5}_{-1.0}$	0.71
	OPT-NIR	0.11 (0.05)	$4.0^{+3.2}_{-1.5}$	0.56
SN2012et	UV-NIR	0.16 (0.02)	$1.8^{+0.6}_{-0.5}$	1.06
	OPT-NIR	0.15 (0.09)	<1.9	1.72
SN2012cp <sup>†</sup>	UV-NIR	0.34 (0.03)	$3.1^{+0.5}_{-0.4}$	0.83
	OPT-NIR	0.36 (0.04)	$3.0^{+0.6}_{-0.5}$	1.03
SN2012bm <sup>†</sup>	UV-NIR	0.45 (0.03)	$3.0^{+0.4}_{-0.3}$	1.68
	OPT-NIR	0.41 (0.05)	$3.5^{+0.6}_{-0.5}$	0.70
SN2012cu	UV-NIR	0.99 (0.03)	$2.8^{+0.1}_{-0.1}$	0.47
	OPT-NIR	1.00 (0.05)	$2.7^{+0.2}_{-0.2}$	0.43
SN2014J	UV-NIR	1.36 (0.02)	$1.4^{+0.1}_{-0.1}$	2.01
	OPT-NIR	1.28 (0.04)	$1.6^{+0.1}_{-0.1}$	2.29

<sup>†</sup>SN colours were compared to SN 2011by

**Table 9.** The best fitted F99 parameters using either only the optical–NIR colours or all colours (UV–NIR) while the  $t_B^{\max}$  and  $s_B$  parameters were fixed to the values given in Tables 5 and 8 respectively. All colour data between phases  $-10$  and  $+35$  days from  $t_B^{\max}$  were used for the fits.

### 6.3 The impact of the UV-data

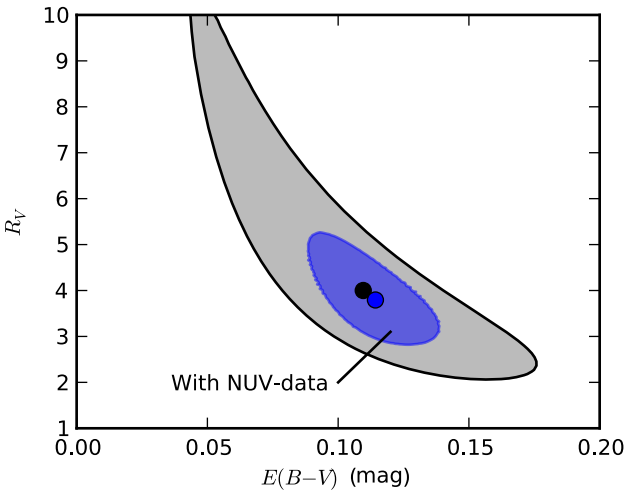
Another approach for testing the robustness of the reddening laws is to study how the fitted parameters are affected by the data at different wavelengths. We do this by comparing the fitted F99 parameter values for two different cases, where we use all measured colours (UV–NIR) and where we omit the UV data (OPT–NIR). A tension between the two cases could be a sign of either assuming an incorrect reddening law or colour reference SN. We keep  $t_B^{\max}$  and  $s_B$  fixed for this study and use SN 2011by as reference for the SNe with red UV colours. The results are presented in Table 9.

The results are consistent for all SNe except for SN 2014J, where the addition of the UV data shifts the fitted values by  $\sim 2$ – $3\sigma$ . We fit a low value of  $R_V$  for this SN and, as shown in Figure 8, a power-law appears to provide a better fit in this case. In Table 10 we present the results on the two colour sets when using a power-law relation for the two low- $R_V$  SNe. While there is a slight,  $\sim 1\sigma$ , tension for SN 2012et, the results are fully consistent between the two cases for SN 2014J.

The major impact of adding the UV-colours is, however, on the uncertainties of the fitted values, in particular for low-reddening SNe. This is illustrated in Figure 16 where the grey contour shows the 68% confidence region for  $E(B - V)$  and  $R_V$  when the F99 law is fitted to the optical–NIR data of SN 2012cg, while the blue contour shows the uncertainties from the full UV–NIR fit. From the figure we conclude that extending the wavelength range significantly improves the constraints on  $R_V$  despite the high intrinsic dispersion adopted for the *Swift* and *HST* colours.

SN	$\lambda$ -range	$A_V$	$\beta$	$\chi^2/\nu$
SN2012et	UV-NIR	0.30 (0.13)	-1.7 (0.5)	0.63
	OPT-NIR	0.16 (0.11)	-3.3 (1.5)	0.82
SN2014J	UV-NIR	1.97 (0.10)	-2.0 (0.1)	1.40
	OPT-NIR	1.93 (0.10)	-2.1 (0.1)	1.02

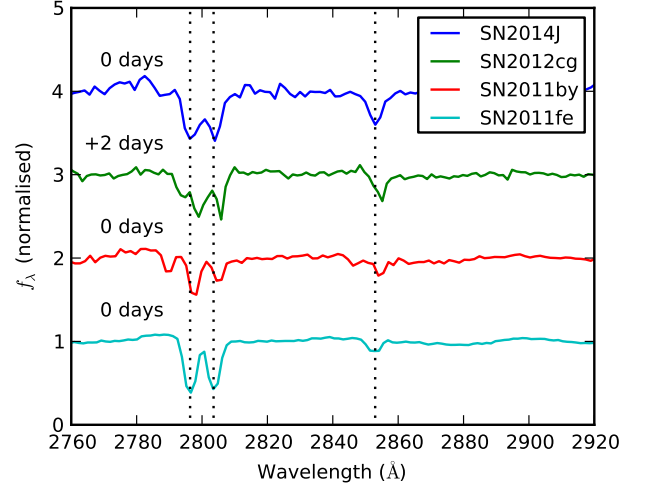
**Table 10.** The best fitted power-law parameters for the two SNe with low  $R_V$  values. The fits were carried out using either only the optical-NIR colours or all colours (UV-NIR) while the  $t_B^{\max}$  and  $s_B$  parameters were fixed to the values given in Table 5. All colour data between phases  $-10$  and  $+35$  days from  $t_B^{\max}$  were used for the fits.



**Figure 16.** Contours showing the joint 68% uncertainty region for the fitted parameters  $R_V$  and  $E(B-V)$  for a F99 extinction laws for SN 2012cg in comparison to SN 2011by. The two contours show the fit results when the full data set (UV-NIR) is used (dark blue contour region) and when the colours based on *HST* and *Swift* UV observations are omitted (light grey contour region). All data between phases  $-10$  and  $+35$  days with respect to  $t_B^{\max}$  were used, and the best fit values for the two cases have been marked by black and blue dots respectively.

#### 6.4 Extinction from interstellar absorption features

In §5 we discussed the constraints on the existence of circumstellar dust that can be derived from non-detection of time-varying Na I D absorption lines. Na I D lines in high-resolution spectra are also commonly used as a proxy for reddening (Poznanski et al. 2011) and the equivalent width of the lines have been shown to correlate well with extinction in the Milky Way (see Munari & Zwitter 1997; Poznanski, Prochaska & Bloom 2012). However, Phillips et al. (2013) found that of their sample of SNe Ia  $\sim 25\%$ , of which most have blue shifted Na I D features, have unusually high Na I column densities compared to the values expected from the Poznanski, Prochaska & Bloom (2012) relation. It is currently unclear whether the high abundance of sodium is intrinsic to the environment of some SN Ia or if the Milky



**Figure 17.** Mg II at  $\lambda\lambda 2796, 2803$  Å and Mg I at  $\lambda\lambda 2853$  Å marked by dotted lines in *STIS* spectra of four SNe with different reddening. The continua of the unsmoothed spectra are normalised and offset for clarity. Above the respective spectra, the phases are indicated at which the spectra were taken.

Way reference value is anomalous compared to other galaxies. Other species, such as diffuse interstellar bands, have been shown to be useful for studying the ISM in SNe Ia host galaxies (Sollerman et al. 2005) and in particular the diffuse interstellar band (DIB) feature at  $5780$  Å may provide a better proxy for reddening of SNe Ia than Na I D (Phillips et al. 2013). In Table 11 we present equivalent width values of Na I D and the DIB at  $5780$  Å with corresponding reddening estimated by the empirical relations (Poznanski, Prochaska & Bloom 2012) and (Phillips et al. 2013), respectively, for the SNe of which we have obtained spectra or published data exists. Further empirical reddening relations exist for ISM features Ca II H&K (Murga et al. 2015) and Mg II (Ménard et al. 2008). These relations have so far not been extensively studied for their validity in estimating the reddening of SNe.

Besides Na I D, narrow Ca II H&K features, with respective equivalent widths of  $770 \pm 30$  and  $579 \pm 35$  mÅ could only be measured from the first FIES epoch of SN 2012cg. These values are well above the range in which the relations in (Murga et al. 2015) are valid. Phillips et al. (2013) also quote the detection of the DIB  $5780$  Å which is consistent with the non-detection in the individual FIES epochs within the signal-to-noise. The most prominent absorption features in Na I D, shown in Figure 10 roughly agree with  $v_h$  from Table 1 and are slightly blueshifted with respect to the apparent stellar velocity of  $\sim 448$  km/s along the line-of-sight of SN 2012cg (see Cortés, Kenney & Hardy 2006). Additional features of the Na I D and Ca II H&K profiles span a range from  $\sim 370$ – $530$  km/s, whereby a number of features are blended on the redshifted side of the profile. SN 2012cg is included in the analysis of Phillips et al. (2013) and is not counted in the group of SNe with anomalously high Na I column densities. However, the best fit  $E(B-V)$  in Table 9 for 2012cg highly disagrees with the value computed from the Poznanski, Prochaska & Bloom (2012) relation.

Four epochs of high-resolution spectra of SN 2012cu

SN	Na I D1 (mÅ)	Na I D2 (mÅ)	$E(B - V)$ (Na I D)	DIB 5780 (mÅ)	$A_V$ (DIB)	Mg II 2796 (Å)	$E(B - V)$ (Mg II)	$E(B - V)$ (Phot.)
2014J	2558 (6)	2831 (9)	-	344 (4)	1.8 (0.9)	$\leq 4.6$ (0.2)	$\leq 0.14$ (0.08)	1.36 (0.02)
2012cg	700 (9)	1018 (8)	$1.6^{+0.4}_{-0.3}$	85 (5) <sup>†</sup>	0.4 (0.2)	$\leq 3.5$ (0.2)	$\leq 0.08$ (0.05)	0.11 (0.02)
2012cu	849 (3) <sup>‡</sup>	925 (3) <sup>‡</sup>	$1.7^{+0.4}_{-0.3}$	-	-	-	-	0.99 (0.03)
2012et	-	650 (40) <sup>*</sup>	$0.3^{+0.3}_{-0.1}$	-	-	-	-	0.16 (0.02)
2011fe	27.4 (1.1) <sup>◊</sup>	47.1 (0.8) <sup>◊</sup>	$0.017^{+0.004}_{-0.003}$	-	-	2.9 (0.2)	$0.06^{+0.03}_{-0.02}$	-
2011by	-	-	-	-	-	$\leq 2.3$ (0.2)	$\leq 0.04$ (0.02)	-

Equivalent width values retrieved from <sup>†</sup> Phillips et al. (2013), <sup>‡</sup> Sternberg et al. (2014), <sup>\*</sup> Maguire et al. (2013) and <sup>◊</sup> Patat et al. (2013).

**Table 11.** Equivalent width values of Na I D, the DIB feature at  $\lambda\lambda 5780$  Å and Mg II at  $\lambda\lambda 2796$  Å for which reddening parameters can be inferred from Poznanski, Prochaska & Bloom (2012), Phillips et al. (2013) and Ménard et al. (2008), respectively. Upper limits are determined for Mg II measurements, since the doublet is blended in the *STIS* spectra. In the case of SN 2014J, the Na I D equivalent width extrapolates the relation by Poznanski, Prochaska & Bloom (2012) beyond any reasonable value for  $E(B - V)$ . Best photometric  $E(B - V)$  values have been taken from Table 9 for comparison and the  $E(B - V)$  for 2011fe and 2011by is assumed to be negligible.

have been described in Sternberg et al. (2014). No time-variability is detected in the deep Na I D absorption features, which likely consist of several blended components. The velocity of the Na I D lines of  $\sim 1130$  km/s are in agreement with the H I velocity along the line-of-sight of SN 2012cu in NGC 4772 (Haynes et al. 2000). We perform the photo-ionisation analysis described in Section 5 on the Na I D doublet of 2012cu. The non-variation of Na I D1 indicates that there is no Na I between  $6 \cdot 10^{18} - 2 \cdot 10^{19}$  cm. However, these constraints are weak since the absorption lines are close to saturation and not optically thin. The  $E(B - V)$  computed from the Poznanski, Prochaska & Bloom (2012) relation is  $\sim 2\sigma$  greater than the photometric value. Furthermore, this value must be considered a lower limit since the Na I D doublet of SN 2012cu appears to be close to saturation.

SN 2012et is included in the sample of Maguire et al. (2013) in which an Na I D2 equivalent width is quoted from a mid-resolution spectrum. The corresponding Poznanski, Prochaska & Bloom (2012) relation  $E(B - V)$  is  $\sim 1\sigma$  greater than expected.

It is clear from Phillips et al. (2013) that Na I D do not serve as good reddening proxies for SNe Ia and that the scatter in the Poznanski, Prochaska & Bloom (2012) relation is underestimated. This appears to be the case for SNe with high reddening, such as 2014J, as well as low reddening such as 2012cg. Furthermore, many of the Na I D profiles do not appear to be optically thin, implying that the column densities are rather underestimated, which would further increase the discrepancy between the Poznanski, Prochaska & Bloom (2012) relation and observations. There is thus motivation to investigate other ISM species, such as e.g. Mg II, which are more difficult to observe and have not been studied in large samples so far.

*STIS* spectra of SNe 2014J (from Foley et al. 2014), 2012cg, 2011fe and 2011by<sup>7</sup> around the Mg II doublet at  $\lambda\lambda 2796, 2803$  Å are shown in Figure 17. The absorption features can be used to test the relation

$$E(B - V) = (0.008 \pm 0.001)W^{(1.88 \pm 0.17)}$$

<sup>7</sup> *STIS* spectra of SNe 2011fe and 2011by were taken as part of Program GO-12298; PI: Ellis

of Ménard et al. (2008), where  $W$  is the equivalent width of Mg II at  $\lambda\lambda 2796$  Å. The *STIS* spectra of SNe 2012bl, 2012bm, 2012cp and 2012et are not included in the sample due to the low signal-to-noise around the Mg II doublet. Upon visual inspection of Figure 17, one can see that the Mg II lines are deep for all four SNe and the line ratios obtained from fitting a double Gaussian profile further indicate that the lines are not optically thin. In all *STIS* spectra, except those of 2011fe, the Mg II doublet is strongly blended. Further, it cannot be excluded that the entire absorption is attributed to the Milky Way. We therefore measure the total equivalent width of the observed features and consider 2/3 of the value to be the upper limit of the component at  $\lambda\lambda 2796$  Å. The fraction is chosen from the line ratio of the Mg II doublet, if it is optically thin. The limits we obtain along with the corresponding  $E(B - V)$  computed from the relation above are shown in Table 11. The Mg II doublet is resolved better in the 2011fe *STIS* spectra and the equivalent width can be determined from the Gaussian fits more accurately. The value presented in Table 11 is the average equivalent width computed from all available 2011fe *STIS* spectra. In comparison to the photometric  $E(B - V)$ , the Ménard et al. (2008) relation clearly underestimates the reddening of 2014J, whereas no firm conclusions can be drawn for the other SNe due to the large errors and since blending with Milky Way features cannot be excluded.

## 6.5 The origin of SN reddening

While it can be expected that the observed SN Ia colour arise from a mixture of extinction by dust and intrinsic SN Ia colour diversity, the latter is typically of order  $|\Delta(B - V)| \lesssim 0.1$  (e.g. Nobili & Goobar 2008). Thus, we can assume that the impact of dimming by dust will dominate with increased observed reddening and that the fitted extinction laws will be less susceptible to intrinsic SN Ia variations. It is therefore particularly interesting to compare the fits for SNe 2012cu and 2014J. They are the reddest objects in our sample, and at the same time, show very different total-to-selective extinction,  $R_V = 2.8 \pm 0.1$  and  $R_V = 1.4 \pm 0.1$  for SNe 2012cu and 2014J respectively.

The evidence for material in the line-of-sight is strength-

ened by the fact that both SNe show deep Na I D absorption lines in their spectra (e.g. Sternberg et al. 2014; Goobar et al. 2014, for 2012cu and 2014J respectively). Since CS dust is expected to give rise to steep extinction laws, such a scenario may offer an explanation for the low  $R_V$  of SN 2014J. However, no time-variability has been found for the Na I D absorption lines for any of the SNe we studied, although Graham et al. (2015b) report a detection of time-varying K I absorption for SN 2014J. When combined with the non-varying Na I D, these findings point to circumstellar dust at radii of  $r_{\text{dust}} \gtrsim 10^{19}$  cm due to the difference in ionisation energies. The detection of light echoes (Crofts 2015) also points to the existence of dust at a comparable distance from the site of the explosion, 11 pc, i.e.,  $r_{\text{dust}} = 3 \cdot 10^{19}$  cm.

Further, we have studied the point spread function shape of SN 2014J in the *HST* A14 photometry to search for signs of dust at smaller radii but none could be detected, which is consistent with the lack of time-evolution of the SN 2014J colours with respect to SN 2011fe. For CS dust at  $r_{\text{dust}} \gtrsim 10^{19}$  cm the photon time-delay exceeds the typical time-scale of the SN Ia light curve and no colour time-evolution is expected as discussed in Amanullah & Goobar (2011).

Other observations suggesting that the nature of the extinction of SN 2014J is primarily interstellar, is that the polarisation angle of the SN light is well aligned with the spiral structure of the host galaxy (Patat et al. 2015), the lack of thermal emission from the SN environment in near and mid-IR (Johansson et al. 2014), and that the velocities of the observed multiple Na I D lines are consistent with measured H I velocities along the line-of-sight in the host galaxy (Ritchey et al. 2015). The latter is also the case for SNe 2012cu (Haynes et al. 2000). Even if dust does exist at  $r_{\text{dust}} \sim 10^{19}$  cm this is not expected to have any major impact on the measured  $R_V$  for such large distances (Wang 2005).

If the reddening of both SNe 2012cu and 2014J are dominated by interstellar extinction, the dust properties along the two lines-of-sight must be significantly different, and the dust in M82 must be of a nature that has not been observed in the Milky Way (although see Brown et al. 2015, for a possible explanation). M82 is a starburst galaxy, and the dust properties have been studied by Hutton et al. (2014). They find that a steep Milky Way-like law is preferred over the Calzetti (2001) law, used for starburst galaxies and they derive a wavelength-dependence of the scattering within a central projected radius of  $< 3$  kpc that suggests that "only small grains are entrained in the inner SNe-driven wind", which is consistent with both the location (see Table 1) and the characteristic of the observed reddening of SN 2014J.

For SNe Ia, Rigault et al. (2013) found that when using H $\alpha$  emission as a tracer for ongoing star formation, SNe with local H $\alpha$  emission is on average redder and more homogeneous, resulting in a lower brightness dispersion. These results have been confirmed by Kelly et al. (2015) and Rigault et al. (2015) using UV data from the *Galaxy Evolution Explorer* (*GALEX*). If the low  $R_V$  originates from small dust grains in regions undergoing intense star formation, we can expect to observe different local properties between SNe 2012cu and 2014J. In Table 12 we present the surface brightness of at the position of the SNe using an aperture with a 2 kpc radius. From this we cannot draw any firm conclusion

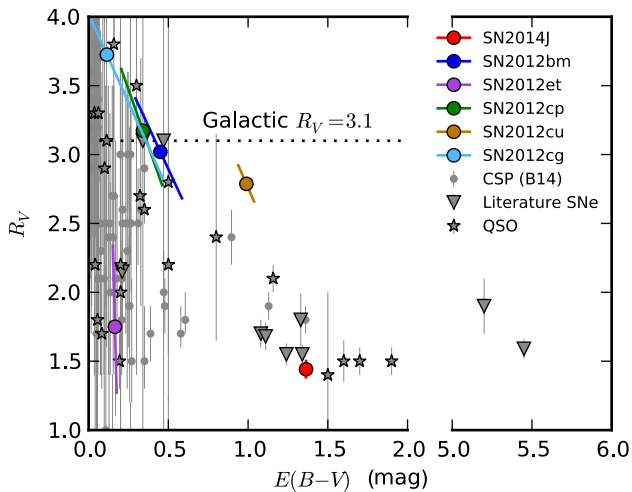
SN	FUV	NUV
2014J	26.5(0.2)	24.9(0.1)
2012bm	26.7(6.1)	25.9(2.4)
2012cg	25.1(0.4)	24.4(0.2)
2012cp	24.7(1.3)	24.6(0.7)
2012cu	26.6(2.3)	26.0(1.0)
2012et	24.9(2.6)	24.4(1.2)
2012bl	26.1(3.3)	26.2(2.1)

**Table 12.** Surface brightness (mag/arcsec) measured using a 2 kpc radius centred on the SN position in the *GALEX* FUV and NUV filters.

based on the local UV surface brightness of the host galaxies. Even after we attempt to correct for host galaxy dust, following the recipe used in Rigault et al. (2015), we see no correlation with UV surface brightness and measured  $R_V$ .

In Figure 18 we show the fitted F99 parameters from Table 5 for all SNe together with similar measurements from the literature. Low values of  $R_V$  are not unique for SNe Ia, and has also been observed in the line-of-sight of other light sources such as gamma-ray burst afterglows (Fynbo et al. 2014) and quasi-stellar objects (QSO, e.g. Fynbo et al. 2013; Leighly et al. 2014). From the figure it is apparent that there are several observations with low  $R_V$  values for high extinction SNe Ia, while SN 2012cu is both showing high extinction and a relatively high  $R_V$ . For low-extinction SNe it is by nature difficult to measure  $R_V$  for individual objects, but there are several statistical studies that obtain higher ( $R_V \lesssim 3$ ) average values of  $R_V$  compared to what is observed for high-extinction SNe Ia (e.g. Mandel, Narayan & Kirshner 2011; Chotard et al. 2011). The UV observations allow us to improve the sensitivity on  $R_V$  for individual low-reddened SNe, and it is interesting to note that we measure  $R_V$  discrepancies similar to what we saw for high-extinction objects. The colours of SNe 2012cg and 2012et are however a mixture of intrinsic variation as was shown by the comparison of SN 2012cg with SNe 2011by and 2011fe in particular. For these objects we therefore cannot draw any firm conclusions to what extent the measured  $R_V$  is solely a dust property. Note though that there are strong indications of dust along the line-of-sight for these SNe from Na I D absorption (as we have seen for SN 2012cg and for SN 2012et in Maguire et al. 2013). Independent of the origin of the reddening we conclude that our sample of reddened SNe Ia with well observed colours, although limited, suggests that a range of  $R_V$  values can be expected for both low and high-reddened SNe Ia.

Given the findings of M13 that all "NUV-blue" SNe belong to the photospheric NV class while the "NUV-red" consists of both NV and HV SNe and the colour-velocity relation for  $B - V$  (Wang et al. 2009), we can also compare the derived extinction parameters with the measured Si II  $\lambda 6355$  Å velocities. From Figures 12 and 18 we conclude that both SNe for which we measure a low  $R_V$  (2014J and 2012et) are classified as HV SNe while the remaining belong to the NV class. This is consistent with the statistical study of samples of SNe Ia by Wang et al. (2009), although our sample size does not allow us to draw firm conclusions. However, the fact that we present SNe with both high and low values of  $R_V$  in both high and low extinction environments

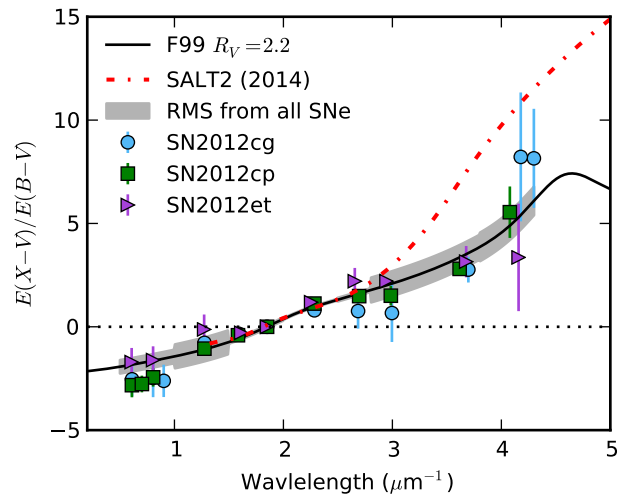


**Figure 18.** Fitted  $R_V$  and  $E(B - V)$  for the full phase range of  $-10$  to  $+35$  days, where SN 2011by has been used as reference for SNe 2012bm, 2012cg, and 2012cp and SN 2011fe was used for the remaining SNe. Also plotted are measurements from the literature, where 9 SNe Ia from Hernandez et al. (2000); Krisciunas et al. (2006, 2007); Benetti et al. (2004); Elias-Rosa et al. (2006, 2008); Hicken et al. (2009) and Folatelli et al. (2010) together with three Type II SNe from Pozzo et al. (2006) and Poznanski et al. (2009) are included. The parameters measured by studying reddening of QSO by foreground galaxies are from Elíasdóttir et al. (2006) and Östman, Goobar & Mörtzell (2006).

argues against that the origin of the discrepancy in observed reddening laws is solely due to intrinsic SN properties.

One possible explanation to both of these findings could be that HV SNe primarily explode in environments with low  $R_V$  dust. Wang et al. (2013) and Pan et al. (2015) have shown that HV SNe are, on average, both located in more luminous hosts, and in brighter regions, closer to their host nuclei, compared to the corresponding NV samples. Further, the results from Hutton et al. (2014) showed that a steep Milky-Way law is preferred for the central region of M82, and similarly the extinction towards the Galactic bulge also follows a steeper extinction curve ( $R_V \approx 2.5$  Nataf et al. 2013) compared to the Milky Way average. Both the dust and SN Ia properties are likely to be affected by the environment where they are produced and if HV SNe Ia originate from younger and more metal-rich progenitors than NV SNe Ia (Wang et al. 2013), it is possible that the average dust properties in the two environments are different as well.

Studies of high-redshift SNe for cosmological applications are unlikely to include highly reddened SNe like 2012cu and 2014J, but objects with similar reddening to SNe 2012cg and 2012et are common. In Figure 19 we show the normalised colour excesses,  $E(X - V)/E(B - V)$ , for the three least reddened SNe, where we have used the approach from Figure 7 to obtain the plotted wavelengths. Figure 19 also shows the average reddening F99 law and the SALT2 colour law (Guy et al. 2007) which has been empirically derived without attempting to discriminate between the different sources of SN reddening. Plotted here is the version used in the cosmology analysis from Betoule et al. (2014). Although the reddening laws are consistent at wavelengths covered by



**Figure 19.** The normalised colour excess  $E(X - V)/E(B - V)$  is shown using the same approach as in Figure 7 for the measurements of the three least reddened SNe. The errorbars are dominated by the adopted intrinsic colour uncertainties. Overplotted is also the average F99 law (solid black) for the three SNe (fitted by comparing the colours to SN 2011fe for SN 2012et and to SN 2011by for the others) and the version of the SALT2 colour law (Guy et al. 2007) used in Betoule et al. (2014). The SALT2 law is only defined over the wavelength range 2000–8000 Å. The grey regions show the RMS of the full sample with respect to the average reddening law for different wavelength bins.

the  $B - R$  bands, the SALT2 law starts to diverge from the data in the UV, wavelengths that becomes increasingly important when SNe Ia at high-redshifts are studied. The lack of understanding SN reddening is one of the dominating astrophysical systematic uncertainties when SNe Ia are used for cosmology (see e.g. Amanullah et al. 2010; Sullivan et al. 2011; Suzuki et al. 2012; Betoule et al. 2014) and further UV studies are needed in order to discriminate between different sources of reddening.

## 7 SUMMARY AND CONCLUSIONS

In this work we have presented and analysed broadband photometry of seven Type Ia supernovae covering the wavelength range  $0.2 - 2 \mu\text{m}$ , obtained mainly from *HST*, *Swift* and the Nordic Optical Telescope.

All SNe are found to be normal with respect to their light curves and optical spectroscopic properties. The SNe have colour excesses in the range  $E(B - V) = 0.0 - 1.4$  mag and we have studied the reddening properties by comparing them to the normal pristine SNe Ia, 2011fe and 2011by. In particular, we have explored three different extinction laws: O’Donnell (1994), Fitzpatrick (1999), and Goobar (2008).

We have further studied the time-evolution of the fitted reddening laws and tested the consistency over a broad wavelength range. For the one SN, 2012cg, we do find indications of time-evolution of the bluest colours and we have investigated if this could be explained by circumstellar dust. Further, we have also searched for time-varying absorption features, in particular Na I D, for SN 2012cg in multi-epoch high-resolution spectroscopy. The high-resolution spectra

were also used to estimate the extinction along the line-of-sight using the absorption of Na I D and Mg II as dust proxies. The validity of using SN 2011fe as a colour template was discussed and compared to the resulting reddening parameters when using SN 2011by.

We summarise the main findings as follows:

(i) We observe a diversity in extinction properties characterised by a total-to-selective extinction,  $R_V$ , ranging from  $R_V \approx 1.4$  to  $R_V \approx 3$ .

(ii) The diversity between the fitted reddening laws for different SNe is significantly greater than the discrepancy between the laws for a single SN. Although, we do find that the empirical F99 law provides a better fit for SNe with  $R_V \sim 3$  while a power-law is preferred for the low- $R_V$  SN 2014J.

(iii) For the two reddest SNe, both with  $E(B - V) > 1$  mag, we observe significantly different values of  $R_V$  with  $R_V = 2.8 \pm 0.1$  and  $R_V = 1.4 \pm 0.1$ , for SN 2012cu and 2014J respectively. Based on their Na I D profiles being consistent with the observed H I velocity fields in their host galaxies, the lack of thermal emission and the observed polarization being aligned with the host galaxy plane for SN 2014J, we conclude that the reddening is likely dominated by extinction in the interstellar medium of the host galaxy. Further, pre-existing circumstellar dust at radii  $r_{\text{dust}} < 10^{19}$  cm is expected to give rise to evolution of the colour excesses,  $E(X - V)$ , but this is not observed for either SN 2012cu nor SN 2014J.

(iv) We compare nebular spectra of SNe 2011fe, 2012cg and 2014J and conclude that all three are very similar at epochs +280 days from maximum, with SNe 2011fe and 2012cg being remarkably similar.

(v) We compare the *STIS* spectra at maximum of SNe 2011by, 2011fe and 2012cg and conclude that while SN 2012cg is well matched to both SNe at  $\lambda > 3800$  Å it is better matched to SN 2011by at shorter wavelengths.

(vi) For SN 2012cg we observe time-evolution in  $E(uvm2 - V)$  with respect to SN 2011fe, but when the object is compared to SN 2011by the evolution is less significant. Circumstellar dust at  $r_{\text{dust}} \sim 10^{17}$  cm is consistent with  $E(uvm2 - V)$  but disfavoured by the lack of evolution of the redder colours. However, this example illustrates that observations over a wide wavelength range provides a powerful tool for separating the origin of colour excess evolution of SNe Ia.

(vii) For low-reddening SNe such as SN 2012cg, adding UV data to optical and near-IR can decrease the statistical uncertainty of  $R_V$  by  $> 50\%$ .

(viii) The two SNe with low  $R_V$  are found to have the highest Si II velocities at maximum which is consistent with previous studies (Wang et al. 2009).

(ix) By measuring the equivalent width from absorption of Na I, we confirm the results of Phillips et al. (2013), and find that the extinction derived for SNe 2012cg, 2012cu and 2012et using the Poznanski, Prochaska & Bloom (2012) relation is higher than what we observe directly for these objects.

Future studies, including a larger sample of low reddening objects covering a wide range in host galaxy morphologies, will be necessary to robustly estimate the impact of the diversity in the wavelength dependence of reddening on the estimates of cosmological parameters from SNe Ia. The cur-

rent work, clearly showing that there is significant range in the extinction parameters, and possibly also extinction laws, should motivate further multi-wavelength studies of SNe Ia, also including the UV.

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## APPENDIX A: SUMMARY OF SNe STUDIED IN THIS WORK

The seven SNe Ia that we study in this work are briefly summarised below:

- **SN 2012bl** was discovered by the Chilean Automatic Supernova Search (CHASE) 42'9 east and 14'6 north of the center of the galaxy ESO 234-19 on Mar. 26.38 UT from an unfiltered image (Pignata et al. 2012). A spectrum (Prieto 2012) was obtained with the 2.5 m du Pont telescope at the Las Campanas Observatory on Mar. 27.28 and object was classified as a SN Ia using SNID (Blondin & Tonry 2007).

- **SN 2012bm**, located 10'95 west and 10'4 north of the centre of the galaxy UGC 8189, was discovered on



Mar. 27th by Puckett et al. (2012). A spectrum was obtained on Mar. 28.05 UT and it was classified (Cappellaro et al. 2012) as a SN Ia using GELATO (Harutyunyan et al. 2008).

- **SN 2012cg** was discovered (Kandrashoff et al. 2012) by the Lick Observatory Supernova Search (LOSS, Filippenko et al. 2001) on May 17.220 UT. The redder than normal SN was located  $17''.3$  east and  $1''.5$  south (Cenko et al. 2012) of the center of the Virgo Cluster member NGC 4424 in a region with many blue stars and disturbed dust lanes. Graur & Maoz (2012) used pre-explosion *HST* images of this well-studied galaxy to rule-out most supergiants as possible binary companions for the progenitor scenario.

- **SN 2012cp** was discovered  $6''.7$  east and  $1''.2$  south of the center of UGC 8713 on May 23.2 UT by the Puckett Observatory Supernova Search (Cox et al. 2012). Milisavljevic (2012) and Zhang, Wu & Wang (2012) obtained spectroscopy of the object on May 25.3 and May 25.7, respectively, and reported that it was consistent with a SN Ia before maximum.

- **SN 2012cu** was discovered on June 14.6 UT by Itagaki et al. (2012)  $3''.1$  east and  $27''.1$  south of the nucleus of the galaxy NGC 4772 and later classified as a SN Ia by Marion et al. (2012). Pre-explosion *Chandra* observations are available of the host galaxy and were used by Nielsen, Voss & Nelemans (2013) to constrain the X-ray emission from the progenitor scenario.

- **SN 2012et**, located  $5''.3$  east and  $0''.8$  north of the center of MCG +04-55-47, was discovered by Rich et al. (2012) on Sep. 12.057 UT, and later classified on Sept. 13.80 UT as a high-velocity SNe Ia (Dennefeld et al. 2012).

- **SN 2014J** was discovered by Fossey et al. (2014) in the nearby galaxy M82 on January 21.805 UT, and later classified by Cao, Kasliwal & McKay (2014). The SN was located  $54''$  west and  $21''$  south of the ill-defined nucleus of the host galaxy.

## APPENDIX B: A SPECTRAL AND COLOUR MODEL FROM SN 2011fe

SN 2011fe is a normal SN Ia with excellent temporal and wavelength coverage with negligible extinction along the line-of-sight and is therefore a suitable object for reddening studies. In A14 we used the available UV-NIR spectroscopic time-series (Pereira et al. 2013; Mazzali et al. 2014) and the NIR light curves (Matheson et al. 2012) to derive the extinction law of SN 2014J. In this work we extend this template with UV photometry (Brown et al. 2012) and optical photometry at 25 to 40 days past *B*-band maximum.

The light curves in the UV F225W, F275W, F336W and the optical *BVRi* bands are first constructed by synthesizing the spectra from Mazzali et al. (2014) and Pereira et al. (2013). A light curve for each filter is then obtained by fitting smoothed splines using *SNooPy* (Burns et al. 2011). For the redder optical *R* and *i* bands, the sparse temporal coverage after 20 days past *B*-maximum does not allow an accurate spline fit of the second bump. We therefore also use the data from Munari et al. (2013) to constrain the spline fit at  $> 25$  days past *B*-max. For the *Swift/UVOT um2* filter and the NIR *JHKs* we fit spline models to the measured light curves from Brown et al. (2012) and Matheson et al. (2012)

respectively. The light curves are then used to calculate the  $(X - V)_0$  colours shown in Figure 5.

As explained in §4 we also need a template of the spectral energy density of the *unreddened* object in order to be able to calculate the expected extinction for each passband for a given extinction model. In A14 the extinction for a given set of reddening law parameters was first calculated for all available spectra of SN 2011fe and the extinction for any given epoch was then obtained using spline interpolations. Similar to the light curve spline model, this approach is adequate for the UV, but due to the sparse spectral coverage, it will cause the model to deviate from the light curve for the redder optical bands and NIR at  $> 25$  days past maximum. In A14 this did not affect the final results due to the high extinction of SN 2014J and the conservative intrinsic colour dispersions adopted in the analysis.

In this work we have developed a more accurate SED template by first resampling the spectral time series used in A14 to a resolution of  $10 \text{ \AA}$ , which are shown as the red spectra in Figure B1. These were then used to calculate the missing wavelength and temporal elements through linear interpolation (shown in grey in Figure B1), and finally the full matrix is mangled to match the light curves described above. The mangling was carried out by first calculating the ratio between the light curve fluxes and the corresponding synthetic SED fluxes. A function was then setup for each phase where the calculated ratios were adopted for the effective wavelengths of the filters. For all wavelengths shorter (longer) than the bluest (reddest) point, the function was assumed to take the same value as in the extreme points, while we interpolate the function value for all intermediate wavelengths, using the calculated ratios. The SED for the given phase was then multiplied by the function. The difference between the synthetic colours calculated from the mangled SED and the spline colour model was always found to be within  $< 0.05$  mag for all epochs, and often even smaller. In Figure B2 these differences are shown for six different colours. Since these differences are within the adopted intrinsic colour uncertainties they will not have any significant impact on the fitted reddening laws.

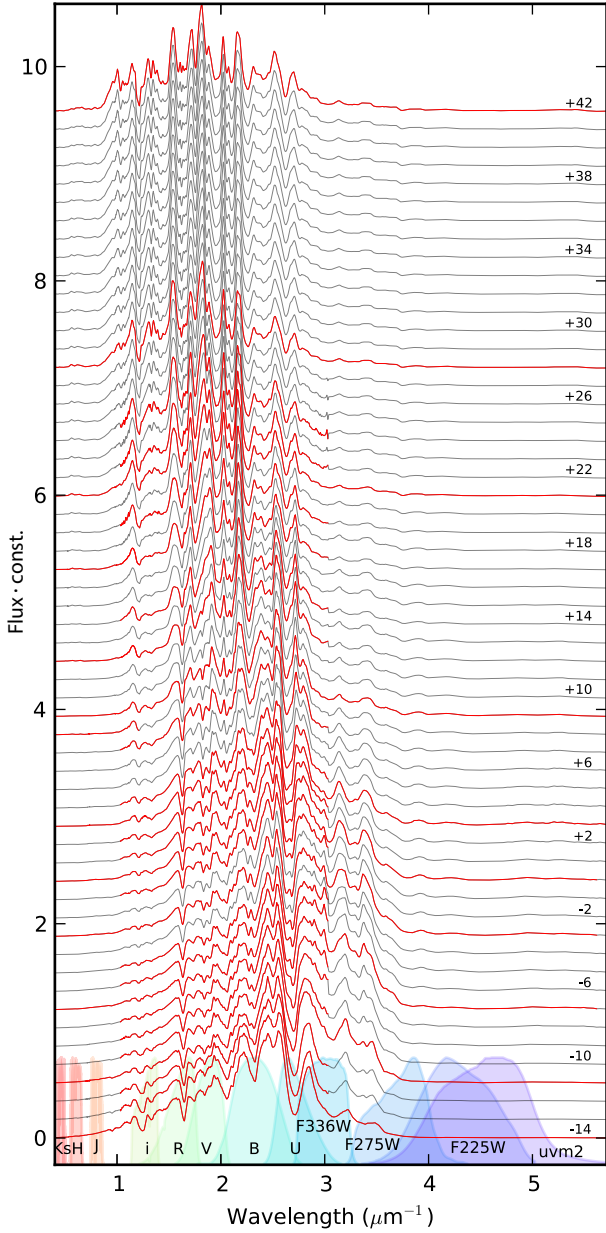
## APPENDIX C: A MODEL FOR THE UV OF SN 2011by

The SNe 2011by and 2011fe have been found to be remarkably similar in the optical, both around maximum and at nebular phases, while showing significant discrepancies in the UV (Foley & Kirshner 2013; Graham et al. 2015a). In Figure 14 the ratio between the spectra are shown at maximum to illustrate this. Unfortunately, SN 2011by does not have the wavelength and temporal coverage that SN 2011fe has, but we can create a crude template for SN 2011by based on the assumption that it is indeed identical to SN 2011fe with the exception of the UV wavelengths.

In Figure 14 the, solid, green line is showing the function

$$f(x') = c \left[ \frac{x'}{(x'^4 + 1)^{1/4}} - 1 \right] + 1 \quad \text{with} \quad x' = \frac{x - x_0}{\alpha} \quad (\text{C1})$$

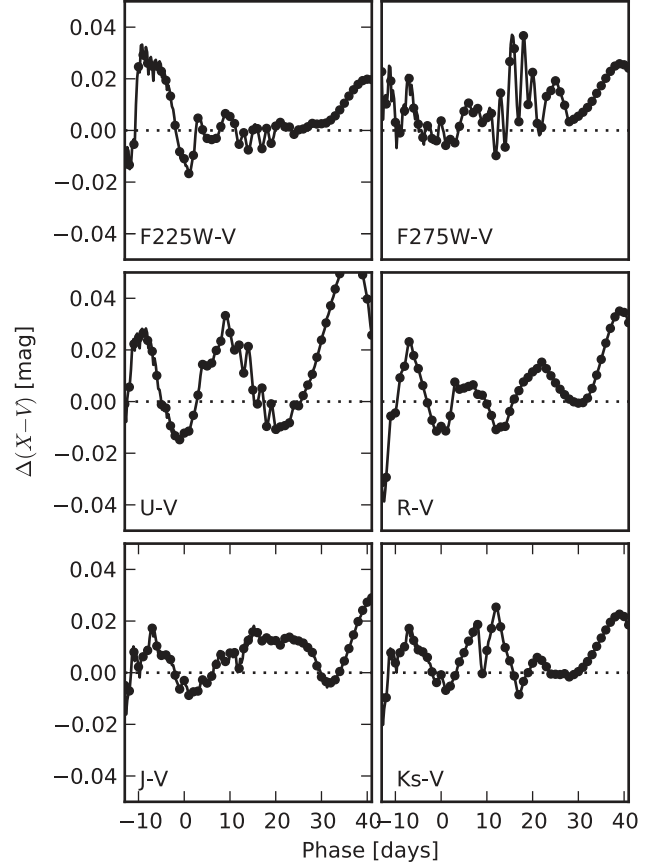
where  $x$  is the wavelength given in microns. In the figure the function is plotted for the parameters  $x_0 = 0.2495$ ,  $\alpha = 0.0360$  and  $c = 0.3359$ .



**Figure B1.** The spectral template constructed from measured spectra of SN 2011fe. Sections based on measured data are shown in red, while the interpolated sections, warped to the light curves are shown in grey. The filters used for the warping are shown at the bottom of the figure.

SN 2011by was observed by *Swift* in the UV (Milne et al. 2013) and in the optical (Silverman, Ganeshalingam & Filippenko 2013), and the colour excesses of these measurements with respect to SN 2011fe are shown in Figure C1. Here we note that  $E(uvm2 - V)$  is showing a similar time-evolution as was seen for SN 2012cg in the upper-left panel of Figure 6.

Under the assumption that SNe 2011by and 2011fe are twins at wavelengths  $\lambda > 3000 \text{ \AA}$  we can create an approximate template for SN 2011by by multiplying the SED of SN 2011fe, presented in Appendix B and shown in Figure B1, by Equation (C1). Further, if we fix the parame-



**Figure B2.** Examples of the difference between synthetic colours calculated from the SED template, and the corresponding colours from the light curve spline templates.

ters  $x_0$  and  $\alpha$  to the values given above, we can fit  $c$  for each epoch with *uvm2* observations by forcing the synthetically calculated  $A_{uvm2}$ , based on the SED to match the observed  $E(uvm2 - V)$  in Figure C1 (again, assuming that  $V^{2011fe} = V^{2011by}$ ). The fitted values of  $c$  are shown in Figure C2 and the synthetic values of  $A_{uvm2}$  that these values correspond to are plotted as black dots in Figure C1.

Further, we parametrise  $c(p)$  with a Gaussian,

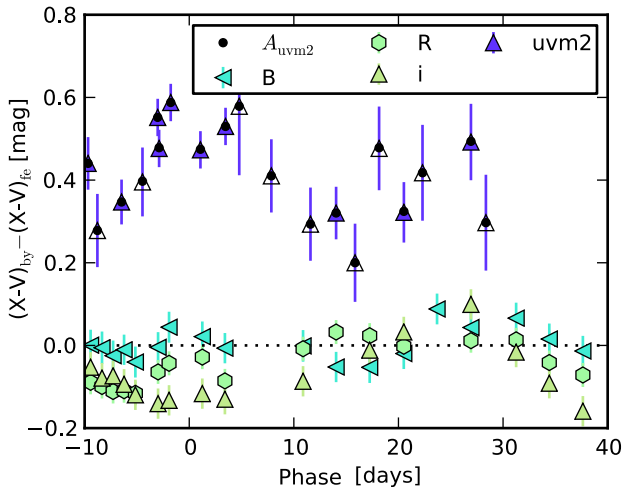
$$c(p) = \frac{a}{\sqrt{2\pi}\sigma} \exp\left[-\frac{1}{2}\left(\frac{p-p_0}{\sigma}\right)^2\right] + b,$$

as shown by the solid line in the Figure C2, with  $p_0 = -0.939$ ,  $a = 1.406$ ,  $b = 0.211$  and  $\sigma = 4.992$ .

To summarise, we create a crude light curve and SED template of SN 2011by by using the template of SN 2011fe from § B and multiply it with Equation (C1), where the parameters, including the time-dependent  $c$ , are given above.

#### APPENDIX D: RED TAILS OF THE *HST* FILTERS

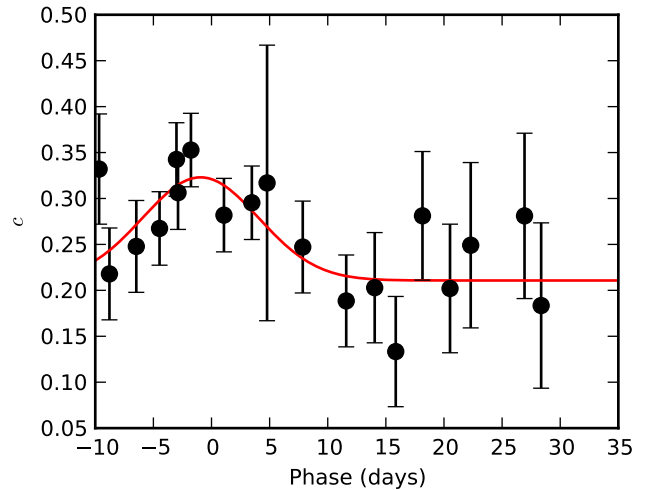
The *HST* filter transmission used in this work, shown in the upper panel of Figure D1, suffer from red tails. That is, part of the flux observed for an object in these filters originates from wavelengths that are redder than the range



**Figure C1.** The difference in colour between SNe 2011by and 2011fe, where the ground based photometry has been adopted from Silverman, Ganeshalingam & Filippenko (2013). Observations for which the colour was obtained from measured  $V$ -band data are shown with filled symbols, while values obtained from a  $V$ -band template are shown with open symbols. The colour difference between the SNe can also be calculated synthetically based on Equation (C1) under the assumption that this function describes the flux ratio of the two objects, for all epochs, after adjusting for the difference in distance. This is motivated by the observation that these SNe appear to be twins at all wavelengths except for the UV. The parameter  $c$  in the equation was fitted by matching the synthetic and measured  $uvrm2 - V$  for each epoch, where the former is shown by the black points for the best fitted values.

where most of the transmission lies. As mentioned in A14, the red tails are relatively small for the  $WFC3/UVIS$  filters, e.g. the transmission of F225W is roughly one part in 100 000 at 5000 Å compared to the peak transmission at  $\sim 2250$  Å. Despite this, the red tails will have implications when reddened SNe Ia are studied. For example, from spectrophotometry of SN 2011fe at maximum (thin, blue line in the middle panel of Figure D1), we conclude that  $\sim 0.5\%$  of the flux (dashed blue line) in the F225W filter will originate from wavelengths  $> 4000$  Å. However, extinction will suppress the UV-flux significantly and after applying a F99 reddening law with  $E(B - V) = 1.4$ ,  $R_V = 1.4$ ,  $\sim 60\%$  of the F225W flux (thick, blue line) comes from photons with  $\lambda > 4000$  Å.

Another way to illustrate the effect of the red tails is shown in the lower panel of Figure D1 where the solid lines show the effective median wavelengths for an observation of SN 2011fe at maximum through the  $WFC$  filters with various amounts of reddening applied. The effective median wavelength is defined such that 50% of the light originates from redder and bluer wavelengths respectively. Further, the wavelength range defined by the points where 10% and 90% of the light comes from bluer wavelengths (i.e. the region that includes 80% of the light) is shown for increasing reddening. From the figure we conclude that, F275W and F336W are less plagued by red tails than F225W, where the latter, in fact, will be probing redder effective wave-



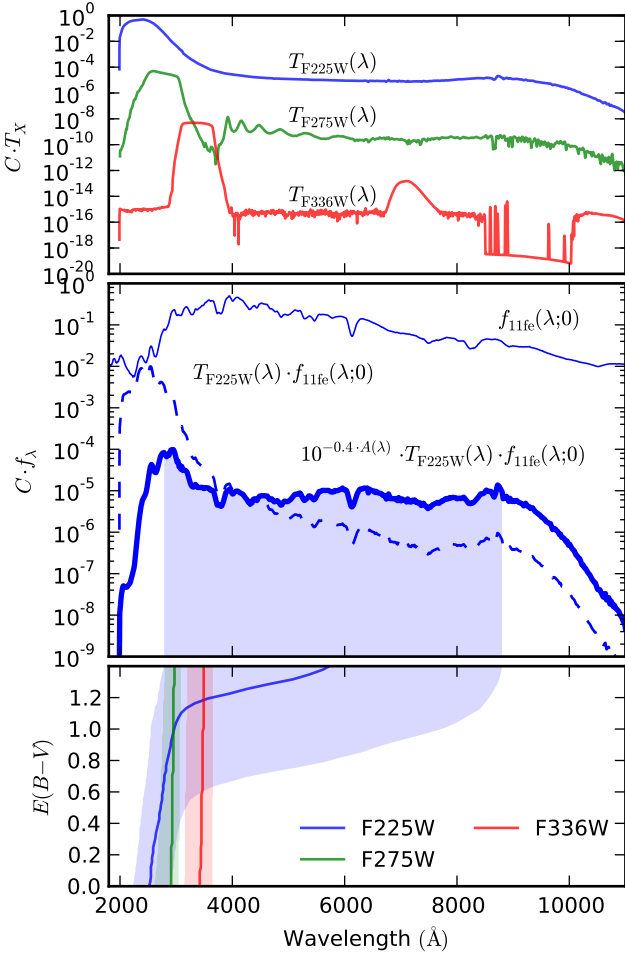
**Figure C2.** The fitted values of parameter  $c$  in after the SED of SN 2011fe has been fitted together with Equation (C1) to match SN 2011by  $uvrm2$  photometry. The solid line is showing a Gaussian fit.

lengths of the SN spectrum than the two former filters for  $E(B - V) \gtrsim 1$  mag.

The fact that we are comparing different effective wavelengths when calculating the colour excess between a reddened and unreddened source is true for all broad band observations. However, fitting extinction properties based on filters with red tails that span a wavelength range where the SED drops rapidly, also puts demands on knowing the passband throughput accurately even when the transmission drops several orders of magnitude as for the filters in Figure D1.

The expected  $WFC3/UVIS$  passband red tails were tested (Brown 2011) in the final round of thermal vacuum tests of WFC3 prior to launch. They studied the filter throughputs in four wavelength bins centered on 450, 600, 750 and 900 nm and found deviations of up to  $\sim 40\%$  between the measured and expected transmission. If the deviations are taken as an estimate of the throughput uncertainty at these wavelengths, we can evaluate its impact on the extinction analysis.

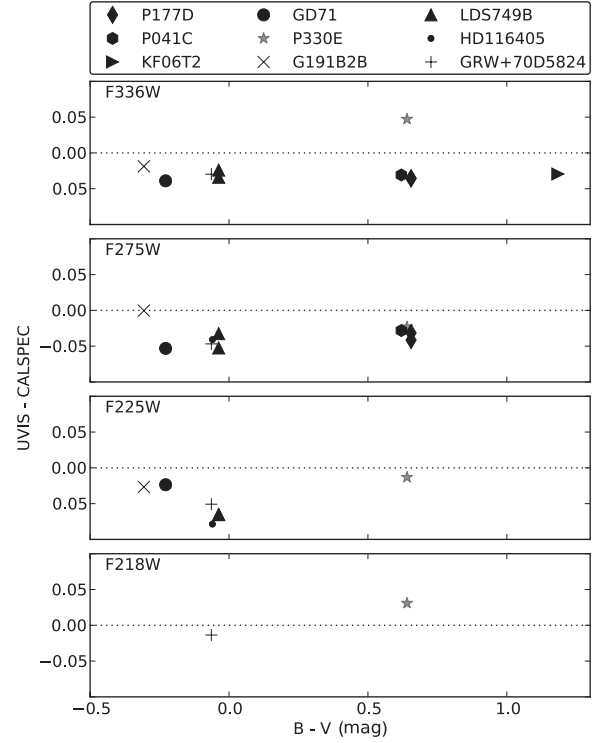
We carried out this test by first applying a F99 reddening law with  $E(B - V) = 1.4$ ,  $R_V = 1.4$  to the SED of SN 2011fe described in Section 4.1. The impact on the calculated extinction for the two bands F225W and F275W (the test of F336W was not carried out due to limitations in the setup as described in Brown 2011) can then be calculated. We do this by comparing synthetic magnitudes from the original SYNPHOT throughputs, to magnitudes obtained after the filters have been multiplied by a spline function derived from the deviations for the four wavelength bins mentioned above. The impact on the calculated extinction was found to always be  $< 0.05$  mag. Further, the worst case scenario can be estimated by letting all the measured deviations from Brown (2011) maximise the amount of red light entering the filter. The result on the extinction from this was  $< 0.10$  mag for  $t < 20$  days with respect to  $B$  max and 0.15–0.20 mag for  $t = 20$ –40 days, which is still within the dispersion that we have adopted for the intrinsic UV colours.



**Figure D1.** The *upper* panel shows the filter transmissions,  $T_X(\lambda)$ , of the *WFC3/UVIS* passbands, multiplied by an arbitrary normalisation,  $C$ . The thin, blue, line in the *middle* panel shows the spectrum of SN 2011fe at *B* band maximum, while the dashed, blue, line shows the same spectrum after it has been multiplied by the F225W throughput. The thick, blue, line shows the same spectrum after it has been reddened with a F99 extinction law with  $E(B - V) = 1.4$  mag and  $R_V = 1.4$ , similar to what we observe for SN 2014J. All three spectra have arbitrary normalisation,  $C$ . In the plot we have also marked the wavelength region from which 80% of the light originates from, defined such that 10% of light comes from bluer wavelengths, and 10% comes from redder wavelengths. The *lower* panel shows the median wavelength, i.e. where 50% of the light originates from redder and bluer wavelengths, for three passbands for SN 2011fe at *B* maximum, reddened with a F99 extinction law ( $R_V = 1.4$ ) with increasing extinction. The 80% region is shown in the same manner as for the middle panel.

It is also possible to investigate the accuracy of the full *HST/WFC3/UVIS* SYNPHOT throughput by comparing spectrophotometry of objects, with known spectra, to *WFC3/UVIS* photometry. This is in fact the most direct test we can carry out of our method for fitting extinction laws since this is using the very same procedure. In figure D2 we have plotted stars from CALSPEC<sup>8</sup> with the corresponding *UVIS*-photometry using the same analysis path as in

<sup>8</sup> <http://www.stsci.edu/hst/observatory/crds/calspec.html>



**Figure D2.** Comparison between spectrophotometry of CALSPEC stars using the *WFC3/UVIS* filter throughputs, with the corresponding *WFC3/UVIS* photometry against the synthetically calculated  $B - V$  colour. The star P330E was used for calibrating the *WFC3/UVIS*.

Section 2.1 with a  $0.4''$  radius. The agreement is in general more than sufficient for the purpose of this analysis, but it is unfortunate that data are not available for redder stars for which the impact of the red tail can be expected to be more significant. The star P330E should be discarded from the comparison since this star has been used for calibrating the *WFC3/UVIS* system.

However, we can carry out the same test for the reddened SNe Ia that have been observed with both *STIS* and *WFC3/UVIS*. In Table D1 we present spectrophotometry at maximum for SNe 2012cg and 2014J, using the transmissions of the *WFC3/UVIS* filters, compared to the corresponding *WFC3/UVIS* photometry. For 2014J the spectrum and photometry were obtained on consecutive days, and the magnitudes based on the former were correct to correspond to the *WFC3/UVIS* epoch using the SN 2011fe template. These corrections are  $< 0.1$  mag.

For these two SNe, the discrepancies between two instruments are within 0.1 mag, including F225W. This is the filter with the most significant red tails, and the photon wavelength distribution for F225W is significantly different for two SNe. The consistency between spectrophotometry and *WFC3/UVIS* photometry for two SNe that define the end-points of the reddening range in our study is showing that any errors of the filter transmissions will not affect the conclusions drawn in the paper.

Further, the fact that we also observe a discrepancy of  $\sim 0.1$  mag for the F336W filter, which has minor red tails,

SN	Filter	<i>STIS</i>	<i>WFC</i>
2012cg	F225W	15.89	15.98 (0.05)
	F275W	13.67	13.61 (0.05)
	F336W	11.85	11.74 (0.06)
2014J	F225W	18.43	18.50 (0.04)
	F275W	16.55	16.55 (0.04)
	F336W	13.15	13.02 (0.04)

**Table D1.** Spectrophotometry of *STIS* spectra at maximum of SNe 2012cg and 2014J at maximum using the filter transmissions shown in Figure D1 compared to *WFC3/UVIS* photometry of the same objects. For SN 2012cg the observations were obtained in consecutive *HST* orbits, while the data for SN 2014J was taken 1 day apart. For this SN, we have calculated the magnitude difference synthetically from the SN 2011fe spectra described in Appendix B. Since the colour evolution is slow between the two observations, any discrepancy in the filter transmissions will be of second order for this correction.

suggests that the discrepancies we observe for F225W may not necessarily originate from uncertainties of these.

## References

- Amanullah R., Goobar A., 2011, *ApJ*, 735, 20  
Amanullah R. et al., 2014, *ApJL*, 788, L21  
Amanullah R. et al., 2010, *ApJ*, 716, 712  
Astier P. et al., 2006, *A&A*, 447, 31  
Benetti S. et al., 2005, *ApJ*, 623, 1011  
Benetti S. et al., 2004, *MNRAS*, 348, 261  
Betoule M. et al., 2014, *A&A*, 568, A22  
Blondin S. et al., 2012, *AJ*, 143, 126  
Blondin S., Prieto J. L., Patat F., Challis P., Hicken M., Kirshner R. P., Matheson T., Modjaz M., 2009, *ApJ*, 693, 207  
Blondin S., Tonry J. L., 2007, *ApJ*, 666, 1024  
Borkowski K. J., Blondin J. M., Reynolds S. P., 2009, *ApJL*, 699, L64  
Bourque M., Anderson J., 2014, CTE Correction for UVIS Subarrays Without Pre-Scan Bias Pixels. Tech. rep., STScI  
Branch D. et al., 2006, *PASP*, 118, 560  
Branch D. et al., 2003, *AJ*, 126, 1489  
Breeveld A. A., Landsman W., Holland S. T., Roming P., Kuin N. P. M., Page M. J., 2011, in American Institute of Physics Conference Series, Vol. 1358, American Institute of Physics Conference Series, McEnery J. E., Racusin J. L., Gehrels N., eds., pp. 373–376  
Brown P. J., 2014, *ApJL*, 796, L18  
Brown P. J., Breeveld A. A., Holland S., Kuin P., Pritchard T., 2014, *Ap&SS*, 354, 89  
Brown P. J. et al., 2012, *ApJ*, 753, 22  
Brown P. J. et al., 2010, *ApJ*, 721, 1608  
Brown P. J. et al., 2015, *ApJ*, 805, 74  
Brown T., 2011, Instrument Science Report WFC3, 2008-49  
Burns C. R. et al., 2014, *ApJ*, 789, 32  
Burns C. R. et al., 2011, *AJ*, 141, 19  
Calzetti D., 2001, *New Astronomy Reviews*, 45, 601  
Cao Y., Kasliwal M. M., McKay A., 2014, The Astronomer’s Telegram, 5786, 1  
Cappellaro E., Pastorello A., Tomasella L., Benetti S., Fiaschi M., Ochner P., Turatto M., Valenti S., 2012, Central Bureau Electronic Telegrams, 3077, 2  
Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, 345, 245  
Cenko S. B., Filippenko A. V., Silverman J. M., Gal-Yam A., Pei L., Nguyen M., Carson D., Barth A. J., 2012, Central Bureau Electronic Telegrams, 3111, 2  
Chomiuk L., Soderberg A., Simon J., Foley R., 2012a, The Astronomer’s Telegram, 4453, 1  
Chomiuk L. et al., 2012b, *ApJ*, 750, 164  
Chotard N. et al., 2011, *A&A*, 529, L4  
Conley A. et al., 2011, *ApJS*, 192, 1  
Conley A. et al., 2008, *ApJ*, 681, 482  
Contreras C. et al., 2010, *AJ*, 139, 519  
Cortés J. R., Kenney J. D. P., Hardy E., 2006, *AJ*, 131, 747  
Cortés J. R., Kenney J. D. P., Hardy E., 2008, *ApJ*, 683, 78  
Cox L. et al., 2012, Central Bureau Electronic Telegrams, 3130, 1  
Crotts A. P. S., 2015, *ApJL*, 804, L37  
Dalcanton J. J. et al., 2009, *ApJS*, 183, 67  
Dennefeld M., Pizzella A., Valentini M., Tomasella L., Benetti S., Cortes M., Ozel N., Rajpurohit A., 2012, Central Bureau Electronic Telegrams, 3226, 2  
Dilday B. et al., 2012, *Science*, 337, 942  
Draine B. T., 2003, *ApJ*, 598, 1017  
Elias-Rosa N. et al., 2006, *MNRAS*, 369, 1880  
Elias-Rosa N. et al., 2008, *MNRAS*, 384, 107  
Elíasdóttir Á., Hjorth J., Toft S., Burud I., Paraficz D., 2006, *ApJS*, 166, 443  
Filippenko A. V., Li W. D., Treffers R. R., Modjaz M., 2001, in Astronomical Society of the Pacific Conference Series, Vol. 246, IAU Colloq. 183: Small Telescope Astronomy on Global Scales, Paczynski B., Chen W.-P., Lemme C., eds., p. 121  
Firth R. E. et al., 2015, *MNRAS*, 446, 3895  
Fitzpatrick E. L., 1999, *PASP*, 111, 63  
Folatelli G. et al., 2013, *ApJ*, 773, 53  
Folatelli G. et al., 2010, *AJ*, 139, 120  
Folatelli G. et al., 2012, *ApJ*, 745, 74  
Foley R. J. et al., 2014, *MNRAS*, 443, 2887  
Foley R. J., Kasen D., 2011, *ApJ*, 729, 55  
Foley R. J., Kirshner R. P., 2013, *ApJL*, 769, L1  
Förster F., González-Gaitán S., Folatelli G., Morrell N., 2013, *ApJ*, 772, 19  
Fossey J., Cooke B., Pollack G., Wilde M., Wright T., 2014, Central Bureau Electronic Telegrams, 3792, 1  
Fox O. D. et al., 2011, *ApJ*, 741, 7  
Fox O. D., Filippenko A. V., 2013, *ApJL*, 772, L6  
Friedman A. S., 2012, PhD thesis, Harvard University  
Friedman A. S. et al., 2014, accepted to *ApJS*, [arXiv:1408.0465](https://arxiv.org/abs/1408.0465)  
Fruchter A. S., et al., 2010, in 2010 Space Telescope Science Institute Calibration Workshop, p. 382–387, pp. 382–387  
Fynbo J. P. U., Krogager J.-K., Venemans B., Noterdaeme P., Vestergaard M., Møller P., Ledoux C., Geier S., 2013, *ApJS*, 204, 6  
Fynbo J. P. U. et al., 2014, *A&A*, 572, A12

- Garavini G. et al., 2007, *A&A*, 470, 411
- Gehrels N. et al., 2004, *ApJ*, 611, 1005
- Goobar A., 2008, *ApJL*, 686, L103
- Goobar A. et al., 2014, *ApJL*, 784, L12
- Goobar A. et al., 2015, *ApJ*, 799, 106
- Goobar A., Leibundgut B., 2011, *Annual Review of Nuclear and Particle Science*, 61, 251
- Graham M. L. et al., 2015a, *MNRAS*, 446, 2073
- Graham M. L. et al., 2015b, *ApJ*, 801, 136
- Graur O., Maoz D., 2012, *The Astronomer's Telegram*, 4226, 1
- Guy J. et al., 2007, *A&A*, 466, 11
- Guy J. et al., 2007, *A&A*, 466, 11
- Hamuy M. et al., 2006, *PASP*, 118, 2
- Harutyunyan A. H. et al., 2008, *A&A*, 488, 383
- Haynes M. P., Jore K. P., Barrett E. A., Broeils A. H., Murray B. M., 2000, *AJ*, 120, 703
- Hernandez M. et al., 2000, *MNRAS*, 319, 223
- Hicken M. et al., 2009, *ApJ*, 700, 331
- Hsiao E. Y., Conley A., Howell D. A., Sullivan M., Pritchett C. J., Carlberg R. G., Nugent P. E., Phillips M. M., 2007, *ApJ*, 663, 1187
- Hsiao E. Y. et al., 2013, *ApJ*, 766, 72
- Hutton S., Ferreras L., Wu K., Kuin P., Breeveld A., Yershov V., Cropper M., Page M., 2014, *MNRAS*, 440, 150
- Itagaki K. et al., 2012, *Central Bureau Electronic Telegrams*, 3146, 1
- Jha S., Riess A. G., Kirshner R. P., 2007, *ApJ*, 659, 122
- Jin Z., Gao X., 2011, *Central Bureau Electronic Telegrams*, 2708, 1
- Johansson J., Amanullah R., Goobar A., 2013, *MNRAS*, 431, L43
- Johansson J. et al., 2014, submitted to *MNRAS*, [arXiv:1411.3332](#)
- Kandrashoff M. et al., 2012, *Central Bureau Electronic Telegrams*, 3111, 1
- Kelly P. L., Filippenko A. V., Burke D. L., Hicken M., Ganeshalingam M., Zheng W., 2015, *Science*, 347, 1459
- Kirshner R. P. et al., 1993, *ApJ*, 415, 589
- Krisciunas K. et al., 2007, *AJ*, 133, 58
- Krisciunas K., Prieto J. L., Garnavich P. M., Riley J.-L. G., Rest A., Stubbs C., McMillan R., 2006, *AJ*, 131, 1639
- Krisciunas K. et al., 2003, *AJ*, 125, 166
- Landolt A. U., 1992, *AJ*, 104, 340
- Leibundgut B., 1988, PhD thesis, PhD thesis. Univ. Basel.137 pp., (1988)
- Leighly K. M., Terndrup D. M., Baron E., Lucy A. B., Dietrich M., Gallagher S. C., 2014, *ApJ*, 788, 123
- Maeda K. et al., 2010a, *Nature*, 466, 82
- Maeda K., Nozawa T., Motohara K., 2014, submitted to *MNRAS*, [arXiv:1411.3778](#)
- Maeda K., Taubenberger S., Sollerman J., Mazzali P. A., Leloudas G., Nomoto K., Motohara K., 2010b, *ApJ*, 708, 1703
- Maguire K. et al., 2012, *MNRAS*, 426, 2359
- Maguire K. et al., 2013, *MNRAS*, 436, 222
- Mandel K. S., Foley R. J., Kirshner R. P., 2014, *ApJ*, 797, 75
- Mandel K. S., Narayan G., Kirshner R. P., 2011, *ApJ*, 731, 120
- Marion G. H., Milisavljevic D., Rines K., Wilhelmly S., 2012, *Central Bureau Electronic Telegrams*, 3146, 2
- Marion G. H. et al., 2015, *ApJ*, 798, 39
- Matheson T. et al., 2012, *ApJ*, 754, 19
- Mazzali P. A. et al., 2014, *MNRAS*, 439, 1959
- McClelland C. M., Garnavich P. M., Milne P. A., Shappee B. J., Pogge R. W., 2013, *ApJ*, 767, 119
- Meikle W. P. S., 2000, *MNRAS*, 314, 782
- Ménard B., Nestor D., Turnshek D., Quider A., Richards G., Chelouche D., Rao S., 2008, *MNRAS*, 385, 1053
- Milisavljevic D., 2012, *Central Bureau Electronic Telegrams*, 3130, 3
- Milne P. A., Brown P. J., Roming P. W. A., Bufano F., Gehrels N., 2013, *ApJ*, 779, 23
- Milne P. A. et al., 2010, *ApJ*, 721, 1627
- Munari U., Henden A., Belligoli R., Castellani F., Cherini G., Righetti G. L., Vagnozzi A., 2013, *New Astronomy*, 20, 30
- Munari U., Zwitter T., 1997, *A&A*, 318, 269
- Murga M., Zhu G., Ménard B., Lan T.-W., 2015, submitted to *MNRAS*, [arXiv:1503.02697](#)
- Nataf D. M. et al., 2013, *ApJ*, 769, 88
- Nielsen M. T. B., Voss R., Nelemans G., 2013, *MNRAS*, 435, 187
- Nobili S., Goobar A., 2008, *A&A*, 487, 19
- Nordin J. et al., 2011, *A&A*, 526, A119
- Nugent P., Kim A., Perlmutter S., 2002, *PASP*, 114, 803
- Nugent P. E. et al., 2011, *Nature*, 480, 344
- O'Donnell J. E., 1994, *ApJ*, 422, 158
- Ofek E. O. et al., 2012, *PASP*, 124, 62
- Östman L., Goobar A., Mörtzell E., 2006, *A&A*, 450, 971
- Pan Y.-C., Sullivan M., Maguire K., Gal-Yam A., Hook I. M., Howell D. A., Nugent P. E., Mazzali P. A., 2015, *MNRAS*, 446, 354
- Parnovsky S. L., Parnowski A. S., 2010, *Ap&SS*, 325, 163
- Patat F. et al., 2007, *Science*, 317, 924
- Patat F. et al., 2013, *A&A*, 549, A62
- Patat F., Cox N. L. J., Parrent J., Branch D., 2010, *A&A*, 514, A78
- Patat F. et al., 2015, *A&A*, 577, A53
- Pereira R. et al., 2013, *A&A*, 554, A27
- Perlmutter S. et al., 1999, *ApJ*, 517, 565
- Perlmutter S. et al., 1997, *ApJ*, 483, 565
- Persson S. E., Murphy D. C., Krzeminski W., Roth M., Rieke M. J., 1998, *AJ*, 116, 2475
- Phillips M. M., 1993, *ApJL*, 413, L105
- Phillips M. M. et al., 2013, *ApJ*, 779, 38
- Pignata G. et al., 2012, *Central Bureau Electronic Telegrams*, 3076, 1
- Poznanski D. et al., 2009, *ApJ*, 694, 1067
- Poznanski D., Ganeshalingam M., Silverman J. M., Filippenko A. V., 2011, *MNRAS*, 415, L81
- Poznanski D., Prochaska J. X., Bloom J. S., 2012, *MNRAS*, 426, 1465
- Pozzo M. et al., 2006, *MNRAS*, 368, 1169
- Prieto J. L., 2012, *Central Bureau Electronic Telegrams*, 3076, 2
- Puckett T. et al., 2012, *Central Bureau Electronic Telegrams*, 3077, 1
- Rau A. et al., 2009, *PASP*, 121, 1334
- Rich D. et al., 2012, *Central Bureau Electronic Telegrams*, 3226, 1
- Riess A. G. et al., 1998, *AJ*, 116, 1009



- Riess A. G., Press W. H., Kirshner R. P., 1996, *ApJ*, 473, 88
- Rigault M. et al., 2015, *ApJ*, 802, 20
- Rigault M. et al., 2013, *A&A*, 560, A66
- Ritchey A. M., Welty D. E., Dahlstrom J. A., York D. G., 2015, *ApJ*, 799, 197
- Roming P. W. A. et al., 2005, *SSR*, 120, 95
- Schlafly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525
- Scolnic D. M., Riess A. G., Foley R. J., Rest A., Rodney S. A., Brout D. J., Jones D. O., 2014, *ApJ*, 780, 37
- Silverman J. M., Filippenko A. V., 2012, *MNRAS*, 425, 1917
- Silverman J. M. et al., 2012, *ApJL*, 756, L7
- Silverman J. M., Ganeshalingam M., Filippenko A. V., 2013, *MNRAS*, 430, 1030
- Silverman J. M., Kong J. J., Filippenko A. V., 2012, *MNRAS*, 425, 1819
- Silverman J. M. et al., 2013, *ApJS*, 207, 3
- Simon J. D. et al., 2009, *ApJ*, 702, 1157
- Smith J. A. et al., 2002, *AJ*, 123, 2121
- Soker N., 2014, *MNRAS*, 444, L73
- Sollerman J., Cox N., Mattila S., Ehrenfreund P., Kaper L., Leibundgut B., Lundqvist P., 2005, *A&A*, 429, 559
- Springob C. M., Masters K. L., Haynes M. P., Giovanelli R., Marinoni C., 2009, *ApJS*, 182, 474
- Stanishev V. et al., 2015, submitted to *A&A*, [arXiv:1505.07707](https://arxiv.org/abs/1505.07707)
- Sternberg A. et al., 2011, *Science*, 333, 856
- Sternberg A. et al., 2014, *MNRAS*, 443, 1849
- Stetson P. B., 1987, *PASP*, 99, 191
- Stritzinger M. et al., 2002, *AJ*, 124, 2100
- Sullivan M. et al., 2011, *ApJ*, 737, 102
- Suntzeff N. B., 2000, in *American Institute of Physics Conference Series*, Vol. 522, American Institute of Physics Conference Series, Holt S. S., Zhang W. W., eds., pp. 65–74
- Suzuki N. et al., 2012, *ApJ*, 746, 85
- Taddia F. et al., 2012, *A&A*, 545, L7
- Telting J. H. et al., 2014, *Astronomische Nachrichten*, 335, 41
- Thomas R. C. et al., 2011, *ApJ*, 743, 27
- Tomasella L. et al., 2014, *Astronomische Nachrichten*, 335, 841
- Tripp R., 1998, *A&A*, 331, 815
- Tully R. B., Rizzi L., Shaya E. J., Courtois H. M., Makarov D. I., Jacobs B. A., 2009, *AJ*, 138, 323
- Valenti S. et al., 2011, *MNRAS*, 416, 3138
- Wang L., 2005, *ApJL*, 635, L33
- Wang X. et al., 2009, *ApJL*, 699, L139
- Wang X. et al., 2012, *ApJ*, 749, 126
- Wang X., Wang L., Filippenko A. V., Zhang T., Zhao X., 2013, *Science*, 340, 170
- Zhang T.-M., Wu C.-J., Wang X.-F., 2012, *Central Bureau Electronic Telegrams*, 3130, 2
- <sup>1</sup> *Oskar Klein Centre, Physics Department, Stockholm University, SE 106 91 Stockholm, Sweden*
- <sup>2</sup> *George P. and Cynthia Woods Mitchell Institute for Fundamental Physics & Astronomy, Texas A. & M. University, Department of Physics and Astronomy, 4242 TAMU, College Station, TX 77843, USA*
- <sup>3</sup> *Cahill Center for Astrophysics, California Institute of Technology, Pasadena, CA 91125, USA*
- <sup>4</sup> *Carnegie Observatories, Las Campanas Observatory, Colina El Pino, Casilla 601, La Serena, Chile*
- <sup>5</sup> *Institute of Theoretical Astrophysics, University of Oslo, P.O. Box 1029, Blindern, N-0315 Oslo, Norway*
- <sup>6</sup> *INAF - Osservatorio Astronomico di Padova, vicolo dell'Osservatorio 5, I-35122 Padova, Italy* <sup>7</sup> *Dark Cosmology Centre, Niels Bohr Institute, University of Copenhagen, Juliane Maries Vej 30, 2100 Copenhagen O, Denmark*
- <sup>8</sup> *Instituto de Astrofísica de Andalucía (IAA-CSIC), Glorieta de la Astronomía s/n, E-18008, Granada, Spain*
- <sup>9</sup> *Unidad Asociada Grupo Ciencias Planetarias UPV/EHU-IAA/CSIC, Departamento de Física Aplicada I, E.T.S., Ingeniería, Universidad del País Vasco UPV/EHU, Bilbao, Spain*
- <sup>10</sup> *INAF- Osservatorio Astronomico di Roma via Frascati 33, 00040, Monte Porzio Catone, Roma*
- <sup>11</sup> *Las Cumbres Observatory Global Telescope Network, 6740 Cortona Dr., Suite 102, Goleta, CA 93117, USA*
- <sup>12</sup> *Department of Physics, University of California, Santa Barbara, Broida Hall, Mail Code 9530, Santa Barbara, CA 93106-9530, USA*
- <sup>13</sup> *Department of Physics and Astronomy, Aarhus University, Ny Munkegade 120, DK-8000 Aarhus C, Denmark*
- <sup>14</sup> *Astrophysics Research Centre, School of Mathematics and Physics, Queen's University Belfast, BT7 1NN, UK*
- <sup>15</sup> *Observatories of the Carnegie Institution for Science, 813 Santa Barbara Street, Pasadena, CA 91101, USA*
- <sup>16</sup> *Benozziyo Center for Astrophysics, Faculty of Physics, Weizmann Institute of Science, 76100 Rehovot, Israel*
- <sup>17</sup> *Oskar Klein Centre, Astronomy Department, Stockholm University, SE 106 91 Stockholm, Sweden*
- <sup>18</sup> *Finnish Centre for Astronomy with ESO (FINCA), University of Turku, Väisäläntie 20, FI-21500 Piikkiö, Finland*
- <sup>19</sup> *Department of Astronomy, University of California, Berkeley, CA 94720-3411, USA*
- <sup>20</sup> *Computational Cosmology Center, Lawrence Berkeley National Laboratory, 1 Cyclotron Road, Berkeley, CA 94720, USA*
- <sup>21</sup> *CENTRA - Centro Multidisciplinar de Astrofísica, Instituto Superior Técnico, Av. Rovisco Pais 1, 1049-001 Lisbon, Portugal*
- <sup>22</sup> *School of Physics and Astronomy, University of Southampton, Southampton, SO17 1BJ, UK*
- <sup>23</sup> *Tuorla Observatory, Department of Physics and Astronomy, University of Turku, Väisäläntie 20, FI-21500 Piikkiö, Finland*
- <sup>24</sup> *Department of Astronomy, Faculty of Mathematics, University of Belgrade, Serbia*
- <sup>25</sup> *Department of Earth and Space Sciences, Chalmers University of Technology, Onsala Space Observatory SE 439-92, Onsala, Sweden*

## SUPPORTING INFORMATION

Additional Supporting Information may be found in the online version of this article.

<sup>26</sup>*ASTRON, Postbus 2, 7990 AA Dwingeloo,  
The Netherlands*

<sup>27</sup>*Kapteyn Astronomical Institute, Postbus 800, 9700 AV,  
Groningen, The Netherlands*