

## Evidence for supernova signatures in the spectrum of the late-time bump of the optical afterglow of GRB 021211<sup>★</sup>

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**Abstract.** We present photometric and spectroscopic observations of the gamma-ray burst GRB 021211 obtained during the late stages of its afterglow. The light curve shows a rebrightening occurring  $\sim 25$  days after the GRB. The analysis of a VLT spectrum obtained during the bump (27 days after the GRB) reveals a suggestive resemblance with the spectrum of the prototypical type-Ic SN 1994I, obtained  $\sim 10$  days past maximum light. Particularly we have measured a strong, broad absorption feature at  $3770 \text{ \AA}$ , which we have identified with Ca II blueshifted by  $\sim 14\,400 \text{ km s}^{-1}$ , thus indicating that a supernova (SN) component is indeed powering the “bump” in the afterglow decay. Assuming SN 1994I as a template, the spectroscopic and photometric data together indicate that the SN and GRB explosions were at most separated by a few days. Our results suggest that GRBs might be associated also to standard type-Ic supernovae.

**Key words.** gamma rays: bursts – supernovae

### 1. Introduction

There is growing evidence that long-duration GRBs (i.e. those lasting more than  $\sim 2$  s; e.g. Fishman et al. 1994) are associated with the death of massive stars. The most compelling case for the existence of a SN-GRB connection is represented by the spatial and temporal coincidence between GRB 980425 and SN 1998bw (Galama et al. 1998). Very recently, in the spectrum of the nearby ( $z = 0.1687$ ) GRB 030329, Stanek et al. (2003) found supernova features, emerging out of the afterglow spectrum and resembling those of SN 1998bw (Patat et al. 2001). Several other possible SN-GRB associations have been suggested (e.g. Wang & Wheeler 1998; Woosley et al. 1998; Germany et al. 2000; Terlevich et al. 1999; Rigon et al. 2003),

although none really compelling due to the poor spatial and temporal coincidences.

The remaining evidence for the existence of a GRB-SN connection relies upon the detection of a rebrightening in the afterglow light curves, 20–30 days after the GRB (e.g. Bloom et al. 1999, 2002; Lazzati et al. 2001). These “bumps” have been interpreted as signatures of SNe at maximum, emerging out of the decaying afterglows. However this evidence and, consequently, the interpretation, are based only on a few photometric measurements in the afterglow light curves and on the assumption that the evolution of the underlying SN is similar to SN 1998bw, except for a luminosity rescaling factor (see e.g. Price et al. 2003). Alternative explanations, such as dust echoes (Esin & Blandford 2000), thermal re-emission of the afterglow light (Waxman & Draine 2000), or thermal radiation from a preexisting SN remnant (Dermer 2003) are not yet ruled out. Indirect evidence that at least some long-duration GRBs are associated with the death of massive stars is provided by the

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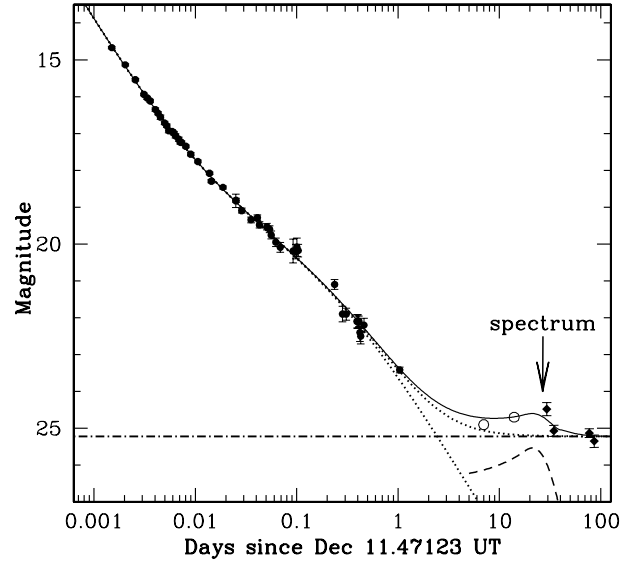
**Table 1.** Summary of our photometric observations. Magnitudes are referred to the complex afterglow + host. Errors are  $1-\sigma$ . \*: average of two measurements obtained on Jan. 9.3 and 10.2.

Start UT	Exp. time	Seeing	<i>R</i> magnitude
2003 Jan. 9.7*	63 min	1.1''	$24.48 \pm 0.18$
2003 Jan. 15.33	15 min	0.8''	$25.07 \pm 0.15$
2003 Feb. 28.02	75 min	0.8''	$25.13 \pm 0.12$
2003 Mar. 9.01	50 min	1.25''	$25.35 \pm 0.17$

detection of star-formation features in the host galaxies (e.g. Djorgovski et al. 1998). Emission and absorption features observed in X-ray spectra (e.g. Piro et al. 2000) are also indicative of large amounts of metals in the GRB surroundings, which can be provided by SN explosions. Finally, the afterglow spectra of GRB 021004 showed absorption systems at different velocities, suggesting the presence of rapidly-moving pre-ejected material surrounding the explosion site (e.g. Møller et al. 2003).

## 2. GRB 021211

GRB 021211 was detected on 2002 Dec. 12 by the HETE-2 satellite at 11:18:34 UT (Crew et al. 2002). The  $\gamma$ -ray fluence was  $(0.96 \pm 0.29) \times 10^{-6} \text{ erg cm}^{-2}$  in the (7–30) keV band and  $(1.98 \pm 0.15) \times 10^{-6} \text{ erg cm}^{-2}$  in the (30–400) keV band (Crew et al. 2003). The GRB was therefore classified as X-ray rich. Given the redshift  $z = 1.006$  of this event (Vreeswijk et al. 2002), the inferred (isotropic) energy was  $E_{\text{iso}} = (6 \pm 0.5) \times 10^{51} \text{ erg}$ , at the low end of the energy distribution of GRBs (Frail et al. 2001). Owing to the rapid distribution of coordinates, an optical afterglow was rapidly discovered by Fox & Price (2002), just 20 min after the GRB onset, as a point-like source with magnitude  $R = 18.29 \pm 0.02$  at coordinates  $\alpha = 08^{\text{h}}08^{\text{m}}59^{\text{s}}.9$ ,  $\delta = +06^{\circ}43'37''.5$  (J2000; Fox et al. 2003). Moreover, the automatic telescopes RAPTOR (Wozniak et al. 2002), KAIT (Li et al. 2003), and Super-LOTIS (Park et al. 2002), imaged the error box a few minutes after the GRB. These observations allowed to monitor the early lightcurve of the afterglow, which could be described by a broken power-law with a flattening at  $t \approx 10 \text{ min}$  (Li et al. 2003). The striking feature of the optical afterglow of GRB 021211 was however its extreme faintness. Compared to other events at similar epochs and redshifts, this afterglow was dimmer by  $\sim 3 \text{ mag}$  in the  $R$ -band (see e.g. Fox et al. 2003). This could in principle be due to heavy extinction within the host. However, Fox et al. (2003) report a broad-band color  $B - K = 3.9$  (on Dec. 11.6 UT), well within the range of GRB afterglows (Šimon et al. 2001). This indicates that GRB 021211 suffered negligible extinction. Moreover, this burst was also underluminous in all observed wavebands: only upper limits were reported in the radio (Fox et al. 2003; Rol & Strom 2002), submillimeter (Fox et al. 2003; Hoge et al. 2002), and TeV (McEnery 2002) regions. Unfortunately, no follow-up X-ray observations were performed.



**Fig. 1.** Light curve of the afterglow of GRB 021211. Filled circles represent data from published or submitted works (Fox et al. 2003; Li et al. 2003; Pandey et al. 2003), open circles are converted from HST measurements (Fruchter et al. 2002), while filled diamonds indicate our data; the arrow shows the epoch of our spectroscopic measurement. The dotted and dot-dashed lines represent the afterglow (see text) and host contribution respectively. The dashed line shows the light curve of SN 1994I reported at  $z = 1.006$  and dereddened with  $A_V = 2$  (from Lee et al. 1995). The solid line shows the sum of the three contributions.

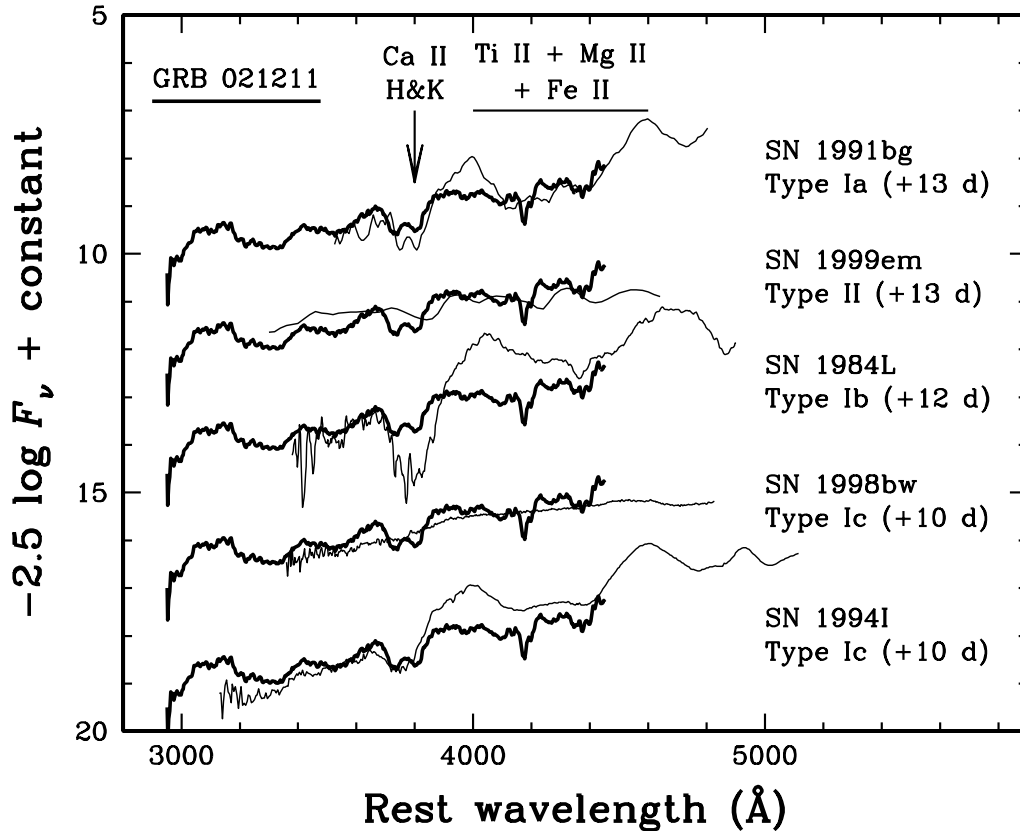
The intrinsic faintness of the optical afterglow made this event a good candidate for showing a prominent late-time bump.

## 3. Photometry

Late-time observations were secured at the ESO VLT-UT4 (Yepun) equipped with the FORS 2 instrument, in the  $R$  band, during the period January–March 2003 (see Table 1).

The conversion to absolute flux was obtained by using both standard calibration and a secondary sequence calibrated, on Feb. 28, with a Landolt standard field (SA98), to account for the observations obtained under non-photometric conditions. Aperture photometry was executed with the packages *apphot* and *photcal* within IRAF, by choosing aperture radii from 0.5'' to 3'', then correcting to infinity with the values found for the standards and using the DAOGROW algorithm. The measurements on the various apertures for the secondary calibrators were found very stable, and we decided to adopt for the target the innermost aperture magnitude, corrected for the aperture.

Our results are listed in Table 1 and supersede our preliminary report (Testa et al. 2003). They have been complemented with a compilation of observations collected from literature and plotted in Fig. 1. A rebrightening is clearly seen, starting  $\sim 15$  days after the burst (Fruchter et al. 2002) and reaching the maximum,  $R \sim 24.5$ , during the first week of January. The contribution of the host galaxy, estimated from our late-epoch



**Fig. 2.** Rest-frame spectrum of the afterglow of GRB 021211 on 2003 Jan. 8.27 UT, or 27 days after the GRB (thick lines), compared with that of several SNe (thin lines).

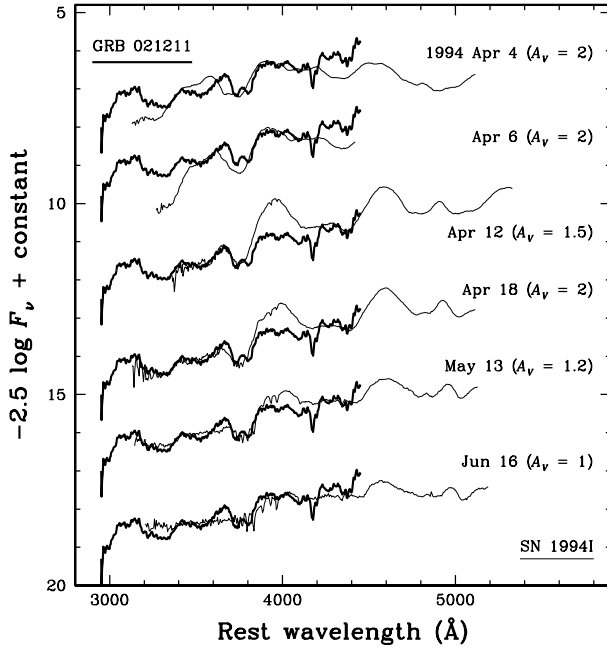
images, is  $R = 25.22 \pm 0.10$ . Therefore, the intrinsic magnitude of the bump was  $R = 25.24 \pm 0.38$ .

The afterglow contribution is more uncertain. The early-time light curve presents several fluctuations, and cannot be easily extrapolated to later epochs. In particular, fitting all data up to 1 day after the GRB<sup>1</sup> with a (convex) broken powerlaw, as used by Li et al. (2003), yields an unacceptable  $\chi^2/\text{d.o.f.} = 115/45$ . We suggest two possible alternatives. First, the shape of the lightcurve is consistent with the presence of two or three rebrightenings underlying a powerlaw component, similarly to what observed in GRB 021004 (Lazzati et al. 2002; Nakar et al. 2003). Alternatively, a second break in the light curve could be present about 10 hours after the GRB. This may be the consequence of the passage of the synchrotron cooling frequency across the optical band (the jet break is expected to occur much later for such a low energy event, according to the correlation of Frail et al. 2001). A fit with a double broken powerlaw (i.e. with *two* breaks, see Fig. 1) yields  $\chi^2/\text{d.o.f.} = 55/43$ ; the improvement is highly significant ( $F$ -test chance probability less than  $10^{-6}$ ). We note that, whichever extrapolation is used, the contribution of the afterglow to the flux measured at the epoch at which our spectrum (the arrow in Fig. 1) was obtained, is negligible (less than 5% in the most conservative case). This fact strongly supports that the bump was powered by a different component other than the afterglow.

#### 4. Spectral analysis

We obtained a spectrum of the afterglow + host with FORS 2, on Jan. 8.27 UT (27 days after the GRB), during the rebrightening phase shown in Fig. 1. The original spectrum covered the range of wavelengths (6000–11 000) Å, although only the interval (6000–9000) Å afforded an acceptable  $S/N$  ( $\geq 3$ ). The resolution was about 19 Å, and the integration time was  $4 \times 1$  h. The slit was rotated in such a way to include also the nearby galaxy reported by Caldwell et al. (2002); this object was well detected and is clearly separated from our target (the seeing in the 4 exposures was  $0''.6$ – $1''.4$ ). The extraction of the spectrum was performed within the MIDAS and IRAF environments, independently by three of us (S.B., M.H., and M.D.V.). We clearly detected the emission line at 7472.9 Å already found by Vreeswijk et al. (2002). Interpreting this as [O II] 3727 Å in the rest frame of the host galaxy, this corresponds to a redshift  $z = 1.006$ . Following this idea, we could also detect, albeit at lower significance, emission lines at 9720 Å, 9955 Å, and 10 025 Å, which we interpret as H $\beta$  and [O III] 4959 Å and 5007 Å. Our spectrum, shown in the rest-frame of the GRB (thick lines in Fig. 2) was smoothed with a boxcar filter (55 Å width) and cleaned from the emission line [O II]. The spectrum of the afterglow is characterized by broad low-amplitude undulations blueward and redward of a broad absorption, the minimum of which is measured at  $\sim 3770$  Å (in the rest frame of the GRB), whereas its blue wing extends up to  $\sim 3650$  Å. We then compared our

<sup>1</sup> We added in quadrature 0.03 mag to the errors of all points in order to account for the use of different telescopes and calibrations.



**Fig. 3.** Comparison of the spectrum of the GRB bump to the ones of SN 1994I taken at different times. Each plot has a different  $A_V$ , in order to better match our spectrum (we adopted the extinction law by Cardelli et al. 1989). For reference, SN 1994I reached its  $B$ -band maximum on 1994 Apr. 9.

spectrum with those of SNe arranged in different spectroscopic types and obtained at different epochs (Fig. 2, thin lines). The comparison includes the type-Ia SN 1991bg (Filippenko et al. 1992; Turatto et al. 1996), the type-II SN 1999em (Hamuy et al. 2001), the type-Ib SN 1984L (Harkness et al. 1987), the peculiar type-Ic SN 1998bw (Galama et al. 1998; Patat et al. 2001), and the type-Ic SN 1994I (Filippenko et al. 1995). Both SN 1999em and SN 1998bw provide a poor match to the afterglow spectrum. Some similarity can be found with the type-Ia SN 1991bg and the type-Ib SN 1984L. A more convincing resemblance is found with the spectrum of the type-Ic SN 1994I, obtained 9 days after its  $B$ -band maximum (Filippenko et al. 1995). The comparison with SN 1994I (and to some extent also with SN 1991bg and SN 1984L) strongly supports the identification of the broad absorption with Ca II H+K; the blueshifts corresponding to the minimum of the absorption and to the edge of the blue wing imply velocities of  $v \sim 14\,400 \text{ km s}^{-1}$  and  $v \sim 23\,000 \text{ km s}^{-1}$  respectively. In principle the Ca II from the host galaxy could contaminate our spectrum. However, since the typical  $FWHM$  of absorption (and emission) lines of galaxies is of the order of  $10\text{--}15 \text{ Å}$  (corresponding to less than  $1000 \text{ km s}^{-1}$ ), and the  $FWHM$  of the observed absorption is about  $150 \text{ Å}$ , one concludes that the Ca II of the host galaxy might affect the observed feature only marginally, if at all.

## 5. Measurements of the time delay between the SN and the GRB

Using SN 1994I as a template, our photometric and spectroscopic data allow us to estimate the time at which the SN exploded, and to compare it with the GRB onset time.

**Spectroscopy.** In order to further study the spectrum of the SN associated with the bump, we have used as a reference (see Fig. 3) the spectral evolution of SN 1994I (Filippenko et al. 1995). To find the best match, different reddenings in the range  $A_V = 1.4 \pm 0.5$  (Richmond et al. 1996) were applied to each spectrum. The epochs before maximum light (April 9) seem to be excluded by the morphology of Ca feature, which is broader and located at shorter wavelengths than in the GRB (the correctness of the wavelength axis was checked by measuring the position of the NaD interstellar absorption in each spectrum). Spectra later than Jun. 1 show a very weak Ca feature and exhibit a poor match with the afterglow spectrum in the  $(3200\text{--}3700) \text{ Å}$  region. Therefore the match with the spectra is acceptable over the range  $\sim$ Apr. 9 to  $\sim$ May 31 ( $\sim 0$  to 40 days after SN maximum light).

Since our spectrum was obtained on Jan. 8, we estimate that the maximum of the SN should have occurred between 2002 Oct. 2 and 2003 Jan. 8 (since the source was at  $z = 1$ ). The exact epoch when the SN exploded depends crucially on the rise time (the time interval from the epoch of the explosion up to maximum light) of type-Ic SNe. The best documented cases are SN 1998bw and SN 1999ex. The latter (Stritzinger et al. 2002) reached  $B$ -band maximum  $\sim 18$  days after the explosion, the former after  $\sim 16$  days (Galama et al. 1998). SN 1994I had a faster rise, reaching its maximum (in  $B$ ) only 12 days after the explosion (Iwamoto et al. 1994). Adopting the latter value, we conclude that the SN exploded between  $\sim$ Sept. 11 and Dec. 15, the later epochs being favored.

**Photometry.** In Fig. 1 we have superimposed to the light curve of the afterglow decay the light curve of SN 1994I (Lee et al. 1995), reported<sup>2</sup> at  $z = 1.006$  and dereddened by  $A_V = 2 \text{ mag}$  (Richmond et al. 1996). The  $K$ -correction has been computed from  $U$ -band data, considering that, at  $z = 1.006$ , the  $U$ -band roughly corresponds to the observed  $R$ -band. The plot (solid line) shows that the luminosity at maximum of SN 1994I ( $M_U = -18.9 \pm 0.3$  assuming  $A_V = 2$ ; see Table 10 of Richmond et al. 1996) agrees very well with that of the bump ( $M_U = -18.8 \pm 0.4$ ). In the figure, a null time delay between the GRB and the SN explosion was used. Letting this delay free to vary did not significantly improve the fit ( $F$ -test chance probability of 36%); the best fit time delay is  $t_{\text{GRB}} - t_{\text{SN}} = (-1.5 \pm 3) \text{ comoving days}$ .

Evidently the photometric observations provide a tighter constraint on the SN-GRB delay. However the uncertainties above are only statistical, while systematics are more difficult to evaluate, especially in consideration of the paucity of observation of type-Ic SNe in their rising phase. Yet, the combination of our photometric and spectroscopic data provide evidence that the SN and the GRB explosions occurred within days from one another, at the most.

## 6. Conclusions

The detection of a broad ( $FWHM \sim 150 \text{ Å}$ ) absorption feature, in the spectrum of the “bump”, which we have identified

<sup>2</sup> Assuming  $\Omega_m = 0.3$ ,  $\Omega_\Lambda = 0.7$ ,  $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$ , the distance modulus at  $z = 1.006$  is  $\mu = 44.07 \text{ mag}$ .



with the Ca II H+K doublet (blueshifted by  $\sim 15\,000\text{ km s}^{-1}$ ) suggests that the rebrightening of the GRB 021211 afterglow was powered by a SN. Assuming for this SN a spectroscopic and photometric behavior similar to that of SN 1994I, our data indicate that the SN and GRB explosion may have occurred almost simultaneously, or at most separated by a few days. The temporal coincidence between the SN and GRB 021211 holds for “short” rise times, of the order of 10–12 days (as observed for SN 1994I; Iwamoto et al. 1994). On the other hand, if a longer rise time (16–20 days, such as that observed in SN 1998bw or SN 1999ex) were used, the conclusion would be that the SN went off several days before GRB 021211.

It is interesting to note that SN 1994I, the spectrum of which provides the best match to that observed in GRB 021211, is a typical type-Ic event rather than an exceptional 1998bw-like object, as the one proposed for association with GRB 980425 and GRB 030329 (Galama et al. 1998; Stanek et al. 2003). If the SN associated with GRB 021211 indeed shared the properties of SN 1994I, this would open the interesting possibility that GRBs may be associated with standard type-Ic SNe, and not only with the more powerful events known as “hypernovae”. This fact may have interesting consequences on the rate of GRB events. One caveat is that the recently studied SN 2002ap (Mazzali et al. 2002) shared some of the properties of hypernovae (e.g. a high expansion velocity), but was not significantly brighter than standard type-Ic SNe. Even if its pre-maximum spectra showed significantly broader lines than our case, this difference vanished after maximum, such that it may not be easy to distinguish between the two types of SNe. However, SN 2002ap had a broader light curve, and it was too faint in the *U*-band (Yoshii et al. 2003). It remains however not firmly established whether GRB 021211 was associated with a standard type-Ic SN or with a “low-energy hypernova” similar to SN 2002ap. We last note that even if GRBs are indeed mainly associated with standard type-Ic SNe, the discovery of overluminous type-Ic events (like SN 1998bw) associated with GRBs is observationally favored, since the SN component can emerge and be observed at early times, when the transient is more frequently monitored.

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## References

- Bloom, J. S., Kulkarni, S. R., Djorgovski, S. G., et al. 1999, *Nature*, 401, 453
- Bloom, J. S., Kulkarni, S. R., Price, P. A., et al. 2002, *ApJ*, 572, L45
- Caldwell, N., McLeod, B., Garnavich, P., & Stanek, K. Z. 2002, *GCN Circ*, 1759
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245
- Crew, G. B., Villaseñor, J., Vanderspek, R., et al. 2002, *GCN Circ*, 1734
- Crew, G. B., Lamb, D. Q., Ricker, G. R., et al. 2003, *ApJ*, submitted [astro-ph/0303470]
- Dermer, C. D. 2003, *ApJL*, submitted [astro-ph/0211300]
- Djorgovski, S. G., Kulkarni, S. R., Bloom, J. S., et al. 1998, *ApJ*, 508, L17
- Esin, A., & Blandford, R. 2000, *ApJ*, 534, L151
- Filippenko, A. V., Richmond, M. W., Branch, D., et al. 1992, *AJ*, 104, 1543
- Filippenko, A. V., Barth, A. J., Matheson, J., et al. 1995, *ApJ*, 450, L11
- Fishman, G. J., Meegan, C. A., Wilson, R. B., et al. 1994, *ApJS*, 92, 229
- Fox, D. W., & Price, P. A. 2002, *GCN Circ*, 1731
- Fox, D. W., Price, P. A., Soderberg, A. M., et al. 2003, *ApJ*, 586, L5
- Frail, D. A., Kulkarni, S. R., Sari, R., et al. 2001, *ApJ*, 562, L55
- Fruchter, A. S., Levan, A., Vreeswijk, P. M., Holland, S. T., & Kouveliotou, C. 2002, *GCN Circ*, 1781
- Galama, T. J., Vreeswijk, P. M., van Paradijs, J., et al. 1998, *Nature*, 395, 670
- Germany, L. M., Reiss, D. J., Sadler, E. M., Schmidt, B. P., & Stubbs, C. W. 2000, *ApJ*, 533, 320
- Hamuy, M., Pinto, P. A., Maza, J., et al. 2001, *ApJ*, 558, 615
- Hamuy, M. 2003, *Review for Core Collapse of Massive Stars*, ed. Fryer (Dordrecht: Kluwer) [astro-ph/0301006]
- Harkness, R. P., Wheeler, J. C., Margon, B., et al. 1987, *ApJ*, 317, 355
- Hoge, J., Willott, C., Grimes, J., Tilanus, R., & Moriarty-Schieven, G. 2002, *GCN Circ*, 1742
- Iwamoto, K., Nomoto, K., Höflich, P., et al. 1994, *ApJ*, 437, L115
- Lazzati, D., Covino, S., Ghisellini, G., et al. 2001, *A&A*, 378, 996
- Lazzati, D., Rossi, E., Covino, S., Ghisellini, G., & Malesani, D. 2002, *A&A*, 396, L5
- Lee, M. G., Kim, E., Kim, S. C., et al. 1995, *J. Korean Astro. Soc.*, 28, 31
- Li, W., Filippenko, A. V., Chornock, R., & Jha, S. 2003, *ApJ*, 586, L9
- Mazzali, P. A., Deng, J., Maeda, K., et al. 2002, *ApJ*, 572, 61
- McEnery, E. 2002, *GCN Circ*, 1740
- Møller, P., Fynbo, J. U., Hjorth, J., et al. 2003, *A&A*, 396, L21
- Nakar, E., Piran, T., & Granot, J. 2003, *New Astron.*, 8, 495
- Park, H. S., Williams, G., & Barthelmy, S. 2002, *GCN Circ*, 1736
- Patat, F., Cappellaro, E., Danziger, J., et al. 2001, *ApJ*, 555, 900
- Piro, L., Garmire, G., Garcia, M., et al. 2000, *Science*, 290, 955
- Pandey, S. B., Anupama, G. C., Sagar, R., et al. 2003, *A&A*, submitted [astro-ph/0304481]
- Price, P. A., Kulkarni, S. R., Berger, E., et al. 2003, *ApJ*, 589, 838
- Richmond, M. W., Van Dyk, S. D., Ho, W., et al. 1996, *AJ*, 111, 327
- Rigon, L., Turatto, M., Benetti, S., et al. 2003, *MNRAS*, 340, 191
- Rol, E., & Strom, R. 2002, *GCN Circ*, 1777
- Šimon V., Hudec R., Pizzichini G., & Masetti N. 2001, *A&A*, 377, 450
- Stanek, K. Z., Matheson, T., Garnavich, P. M., et al. 2003, *ApJL*, in press [astro-ph/030473]
- Stritzinger, M., Hamuy, M., Suntzeff, N. B. et al. 2002, *AJ*, 124, 2100
- Terlevich, R., Fabian, A., & Turatto, M. 1999, *IAU Circ*, 7269
- Testa, V., Fugazza, D., Della Valle, M., et al. 2003, *GCN Circ*, 1821
- Turatto, M., Benetti, S., Cappellaro, E., et al. 1996, *MNRAS*, 283, 1
- Vreeswijk, P. M., Fruchter, A., Hjorth, J., & Kouveliotou, C. 2002a, *GCN Circ*, 1785
- Wang, L., & Wheeler, J. C. 1998, *ApJ*, 504, L87
- Waxman, E., & Draine, B. T. 2000, *ApJ*, 537, 796
- Woosley, S. E., Eastman, R. G., & Schmidt, B. P. 1998, *ApJ*, 516, 788
- Wozniak, P., Vestrand, W. T., Starr, D., et al. 2002, *GCN Circ*, 1757
- Yoshii, Y., Tomita, H., Kobayashi, Y., et al. 2003, *ApJ*, in press [astro-ph/0304010]