

Exotic holographic RG flows at finite temperature

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ABSTRACT: Black hole solutions and their thermodynamics are studied in Einstein-scalar theories. The associated zero-temperature solutions are non-trivial holographic RG flows. These include solutions which skip intermediate extrema of the bulk scalar potential or feature an inversion of the direction of the flow of the coupling (bounces). At finite temperature, a complex set of branches of black hole solutions is found. In some cases, first order phase transitions are found between the black-hole branches. In other cases, black hole solutions are found to exist even for boundary conditions which did not allow a zero-temperature vacuum flow. Finite-temperature solutions driven solely by the vacuum expectation value of a perturbing operator (zero source) are found and studied. Such solutions exist generically (i.e. with no special tuning of the potential) in theories in which the vacuum flows feature bounces. It is found that they exhibit conformal thermodynamics.

KEYWORDS: AdS-CFT Correspondence, Black Holes, Gauge-gravity correspondence

ARXIV EPRINT: [1805.01769](https://arxiv.org/abs/1805.01769)

¹<http://hep.physics.uoc.gr/~kiritsis/>.

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1 Introduction and summary

The gauge/gravity duality [1–3] gives a geometric representation of the renormalisation group (RG): the RG flow in a d -dimensional field theory is realised as the radial flow of d -dimensional hyper-surfaces which foliate a solution of a higher-dimensional (super)gravity theory [4–11]. When the $(d+1)$ -dimensional solution is driven by scalar fields, and when the d -dimensional radial slices are flat, holographic RG flows with a regular interior geometry connect two extrema of the scalar potential $V(\phi)$, where the geometry approaches AdS_{d+1} . Usually, a maximum of V maps to the field theory UV, while a minimum corresponds to the IR (although in special cases there may be exceptions to this rule, [12]).

In this context, the holographic RG flow often captures non-perturbative features of the dual field theory: phenomena such as IR fixed points, confinement, and the condensation of scalar operators, which are usually inaccessible in perturbation theory around a UV conformal fixed point, can be realised in holographic RG flow solutions of simple 2-derivative gravitational theories.

Besides reproducing holographic versions of the field-theoretical non-perturbative features mentioned above, holographic RG flows have been shown to display some unusual, or *exotic*, phenomena. Recently, a systematic exploration of different classes of these exotic features has been initiated in [12] for single scalar field models. This work focused on the analysis of asymptotically AdS holographic RG-flow solutions in $d + 1$ -dimensional Einstein gravity coupled to a single scalar field, dual to renormalisation group flows of a d -dimensional CFT deformed away from the UV by a single operator. This analysis was extended in [13] to multi-field models.

Depending on the details of the bulk scalar potential $V(\phi)$, some of the following exotic feature may arise:

1. *Multiple vacuum flows.* If $V(\phi)$ displays several maxima and minima, there may be several regular flows connecting the same UV maximum to different IR minima. This corresponds to the existence, in the same UV-deformed CFT, of multiple vacua which are distinguished (in the UV) by the value of the scalar condensate $\langle O \rangle$ and which reach different IR fixed points. Only one of them (the one with lowest free energy) is the true vacuum.
2. *Bouncing RG flows.* These correspond to solutions where the flow of the scalar field is non-monotonic. For example, after starting off away from the UV fixed point in the positive direction, the scalar field reaches a maximum (*bounce*), start decreasing

again, and eventually reaches an IR fixed point situated on the opposite side of the initial UV fixed point. At the turning point, the holographic β -function has a square-root branch singularity.

3. *Irrelevant vev flows.* Usually, holographic RG flows with a regular interior correspond to solutions which asymptote in the UV to a maximum of the scalar potential. However in special cases, if the scalar potential is appropriately tuned, regular solutions may exist which connect two minima of the bulk potential (one in the UV, one in the IR). Although they do not exist generically, these flows are interesting because they correspond to a deformation of the theory by the vev of an irrelevant operator (for which turning on a source is not an option). These flows have a continuous moduli space, parametrised by the vacuum expectation value (vev) of the condensate $\langle O \rangle$, with degenerate free energy (also degenerate with the fixed-point AdS solutions where $\langle O \rangle = 0$). For $\langle O \rangle \neq 0$, these solutions break conformal invariance spontaneously and the presence of a moduli space indicates that there is a massless dilaton in the spectrum of excitations, [14, 15].

The discussion above refers to vacuum (i.e. Poincaré invariant) solutions. Probing these features by turning on additional sources, or by considering non-vacuum states, is important in order to test the robustness and consistency of these solutions, most of which arise in bottom-up holographic models. A first step in this direction was taken in [16], where solutions were considered in which the dual field theory lives on a maximally symmetric curved manifold: in this case, the relevant gravity dual solutions are holographic RG flows in which each slice transverse to the (holographic) radial coordinate is a maximally symmetric curved space-time.

Turning on non-zero curvature introduces a new source on the boundary CFT, and correspondingly a new scale beyond the (dimensionfull) coupling of the relevant CFT deformation. As shown in [16], this leads to several interesting new effects. For theories admitting multiple vacuum RG flows (point 1 above) a curvature-driven quantum phase transition is found. Also, certain CFT deformations which do not correspond to any regular vacuum solution become allowed if a sufficiently large curvature is turned-on.

A different way to probe these models is by turning on finite temperature and this is the subject of the present work. Specifically, we consider black-hole solutions of the bulk Einstein-scalar theories, like those discussed in [12], which allow “exotic” vacuum RG-flow solutions.¹

In the same way as space-time curvature, a finite temperature introduces a new scale in the system. Indeed, many of the results we describe here (space of solutions, phase diagrams) closely resemble those found in [16], if we simply replace curvature by temperature as a control parameter.

However there is a very important difference: while a non-zero curvature is a modification of the theory itself, i.e. it introduces new sources (in this case non-trivial components

¹A specific example of the black hole solution obtained by finite T extensions of these vacuum RG flows were first studied in [17, 18].

of the metric, which turns on a relevant deformation proportional to the stress tensor), going to finite temperature corresponds to considering different *states* of the same theory whose vacuum state is a Poincaré-invariant solution.

Indeed, one of the main motivations for considering states beyond the vacuum is to probe whether the theories displaying exotic behaviours discussed above (and in particular the “bouncing” geometries) are “good” holographic theories. From the analysis in [12], no pathology emerged neither in the vacuum (e.g. singular nature of some curvature invariant), nor in the spectrum of small fluctuations around the vacuum, which was shown in general to be perturbatively stable. We note however there may be dynamical instabilities found in the corresponding finite temperature black-hole extensions of these solutions, as was first observed in [17, 18].

By going to finite temperature, we will probe different states of the theory in a way which is complementary to turning on small excitations above the vacuum. In this way, we will test the consistency from the thermodynamic standpoint. As we will see, this will enable us to detect, in some cases, certain pathologies which indicate that some of the holographic models under investigation may not be self-consistent after all.

A particularly interesting question concerns the fate of the irrelevant operator flows discussed in point 3 above. As we will see, at $T \neq 0$ no regular black hole solutions with the same UV asymptotics at the minimum of the potential can be found. This does not come as a surprise as we expect that turning on finite temperature in a theory with a moduli space, lifts generically the moduli space and the moduli acquire a non-trivial potential. This is the case, for example, in the Coulomb branch of the original $N = 4$ AdS/CFT duality.

However, surprisingly, we find that finite temperature allows new irrelevant vev flows from a minimum to exist for *generic* (i.e. non-finely tuned) potentials. These solutions display interesting properties such as a conformal thermodynamics and temperature-driven operator condensation and they may also have interesting hydrodynamic properties.

1.1 Summary of results

In this work we consider $(d + 1)$ -dimensional Einstein gravity coupled to a scalar field ϕ . The action contains a scalar potential $V(\phi)$ with several AdS extrema, which play the role of UV and IR fixed points for asymptotically AdS RG-flow solutions. These are dual to RG flows driven by deforming a UV CFT by a relevant operator O , of dimension $\Delta < d$. Its coupling j is encoded in the leading term of the UV asymptotics of the scalar field about the fixed point value ϕ_{UV} ,

$$\phi(u) \simeq \phi_{UV} + j \ell^{(d-\Delta)} \exp[(d - \Delta)u/\ell], \quad u \rightarrow -\infty, \tag{1.1}$$

where $\ell^2 = -d(d - 1)/V(\phi_{UV})$ is the squared UV AdS length.

We consider the theory at finite temperature by looking at static, planar black hole solutions, of the form

$$ds^2 = \frac{du^2}{f(u)} + e^{A(u)} [-f(u)dt^2 + \delta_{ij}dx^i dx^j], \quad \phi = \phi(u) \tag{1.2}$$

where u is the holographic radial coordinate and (t, x^i) are the space-time coordinates of the dual field theory, whose temperature equals the Hawking temperature of the black hole.

The main goal of this paper is to analyse the space of solutions of the form (1.2) and study their thermodynamics in the canonical ensemble, i.e. at fixed temperature T and fixed UV coupling j .

The main features of the solutions depend only on the dimensionless parameter $\mathcal{T} = T|j|^{-1/(d-\Delta)}$. Requiring the solution to be regular, both \mathcal{T} and the dimensionless vev parameter $C \equiv \langle O \rangle |j|^{-\Delta/(d-\Delta)}$ are completely determined by the value the scalar field takes at the horizon, where $f(u) = 0$. Therefore, a useful way to scan the space of black hole solutions is by changing the horizon endpoint ϕ_h of the flow, and studying the behaviour of the functions $\mathcal{T}(\phi_h)$ and $C(\phi_h)$. The free energy is then computed by standard holographic techniques. We should note that ϕ_h parametrises in a faithful way the black-hole solutions, as there is at most a single solution for each value of ϕ_h .

After a brief general discussion of the thermodynamics of holographic RG-flows in section 2, we study black-hole solutions and their thermodynamics in two specific models:

- a) The first one admits two distinct regular vacuum RG-flows from the same UV theory to two different IR fixed points (section 3).
- b) Next we turn to a models for which the vacuum RG-flow presents a bounce (section 4).

Finally, in section 5 we consider the special black-hole solutions with $j = 0$, which correspond to black holes driven only by the vev of the deforming operator, and have special properties.

In the rest of this subsection we briefly summarise our results.

First order transitions in multi-vacuum theories. The first model we consider has a potential shown schematically in figure 1. For $T = 0$, there are two regular solutions connecting the fixed point UV_1 situated at the origin with each of the two IR fixed point at the two minima. In [12] it was shown that the favoured solution is the *skipping* one, i.e. the one that does not stop at IR_1 but reaches the next available fixed point IR_2 . At finite, fixed temperature T we find there are up to three competing black-hole solutions, all with UV asymptotics at the origin. Two of them skip IR_1 , but exist only up to a maximal temperature T_{\max} . Calculation of the free energy shows a first order phase transition at $T_c < T_{\max}$ above which the *non-skipping* solution becomes dominant. This solution continuously connect to the zero temperature flow ending at IR_1 . Therefore, in this model there is a transition between skipping and non-skipping behaviour as the temperature is increased above T_c .

Thermal desingularisation of ill-defined vacua. We next consider the same potential shown in figure 1, but focusing on the black holes which asymptote in the UV to the fixed point UV_2 . For $T = 0$, there is only one regular RG flow going from UV_2 to IR_1 (with $j < 0$), but no solution with $j > 0$ reaching IR_2 from UV_2 . From the dual field theory standpoint, this means that we can only deform the UV_2 CFT for $j < 0$, but there is no

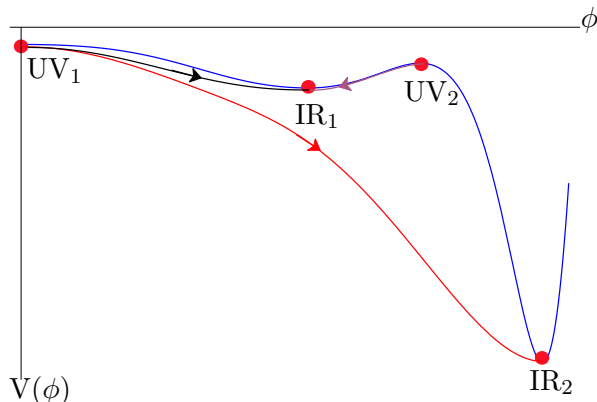


Figure 1. Schematic plot of the potential which allows skipping solutions, with a representation of the corresponding zero-temperature RG flows connecting different fixed points.

well-defined vacuum with $j > 0$. This is not uncommon in perturbative field theory, i.e. pure YM theory exists only for $g^2 > 0$ or $\lambda\phi^4$ theory exists only for $\lambda > 0$.

At finite temperature the situation is richer:

1. For $j < 0$ we find two black-hole solutions at all temperatures, one of which displays a bounce (i.e. an inversion in the flow of the scalar field). Calculation of the free energy shows that the dominant solution is the non-bouncing one at all temperatures
2. For $j > 0$, although there was no regular solution at $T = 0$, two black-hole solutions appear *above a minimal temperature* T_{\min} . This implies that, at high enough temperature, the theory with $j > 0$ may be well-defined, though the zero-temperature vacuum did not exist.

Theories with bouncing vacua and their thermodynamic (in)consistency. The second example we analyse in detail is a potential for which, at zero temperature, the RG-flow solution *bounces*: the scalar field starts decreasing away from the UV, reaches a minimum, then it increases again past the starting point to reach an IR fixed point on the opposite side (see figure 2).

For the potential considered in section 4, the flow represented in figure 2 is the only regular solution for $T = 0$ and $j < 0$ (for $j > 0$ we have its mirror solution with $\phi \rightarrow -\phi$, as the potential we are considering is an even function of ϕ). At finite temperature, for fixed j , we find up to *five* different black-hole branches. One of them connects to the vacuum $T = 0$ solution, one connects with the AdS-Schwarzschild black hole at the origin as $T \rightarrow +\infty$. At low temperature, all solutions exhibit a bounce as in figure 2, while above a certain temperature, new solutions appear which do not bounce, but have horizon (for $j < 0$) on the negative side of the UV fixed point.

The computation of the free energy reveals a puzzling situation. While at high temperature the dominant solution is, as one may have suspected, the large black hole whose horizon is closest to the UV fixed point, the transition to this solution is discontinuous: the free energy shows a jump from the bouncing to the non-bouncing solution as soon as the

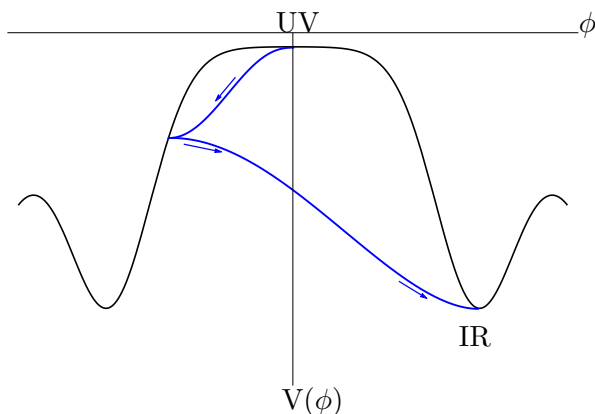


Figure 2. Sketch of a bouncing RG-flow at $T = 0$, with $j < 0$.

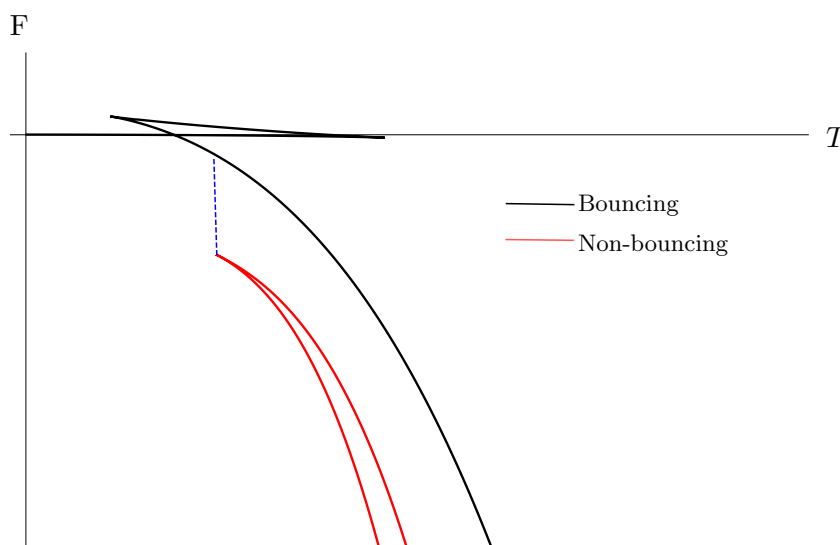


Figure 3. Phase diagram of $j < 0$ black holes in the model with bouncing vacuum solution. Bouncing black holes dominate at low temperature, then the free energy F jumps by a finite amount (dashed blue line) on the non-bouncing branch.

latter one appears (see figure 3). This situation does not allow for a consistent Maxwell construction of the phase diagram, and it may indicate that this is not a good holographic theory, albeit the vacuum was found in [12] to be perturbatively stable *and* the dominant branch at finite T is also thermodynamically stable (the specific heat is positive). Other options are also possible, as we discuss in section 1.2, and at this stage we cannot determine with certainty the reason behind the unusual behaviour.

Vev-driven black holes. When the black-hole horizon ϕ_h approaches a UV or an IR fixed point, this corresponds to a high-temperature or low-temperature limit, respectively. However, in the space of solutions, new infinite-temperature and zero-temperature limits appear which are not connect to any fixed point solution and for which the horizon is far from the extrema of V . These limits signal the existence of *vev-driven black holes* for which, at fixed $T \neq 0$, the source of the deforming operator is set to zero.

These solutions arise in two different situations:

1. At the interface between bouncing and non-bouncing black holes branches asymptoting to the same UV maximum of $V(\phi)$, since at this interface the source j changes sign;
2. When the bounce in a solution coincides with a minimum of the potential.

In the latter case, the $j = 0$ solution can be shown to asymptote in the UV to a *minimum* of the potential rather than a maximum. The corresponding deformation is driven by an *irrelevant* operator, unlike the case in which the UV fixed point is a maximum of $V(\phi)$. Vev-driven black holes are isolated in the space of solutions in the following sense: they correspond to special values of ϕ_h at which the dimensionless temperature $\mathcal{T}(\phi_h) \rightarrow 0, +\infty$.

Vev-driven flows have a very simple thermodynamics, which turns out to be exactly conformal, with the free energy given, for all $T > 0$, by

$$\mathcal{F} = -\sigma T^d,$$

where σ is a temperature-independent coefficient. In addition, the vev of the operator dual to ϕ is completely determined by the temperature,

$$\langle O \rangle = cT^\Delta,$$

where c is another constant.

Finally, we analysed a particular case (with a specially tuned potential) where a regular vev-driven flow solution does exist for $T = 0$, and was constructed in [12]. The corresponding vacuum flow connects two minima of V and it provides a zero-temperature example of a regular flow driven by the vev of an irrelevant operator. We find that in this case, after turning on temperature, one cannot find any regular black-hole solution in which the scalar has a non-trivial flow.

1.2 Discussion and outlook

The results described in the previous section show that the space of black-hole solutions, built around holographic RG flows may have an extremely rich structure and may display unexpected phenomena.

The structure of the different branches of solutions at finite temperature closely parallel similar structures that were found in [16] when the dual field theory is defined on a positively curved space-time. Furthermore, as we will discuss below, both curvature and temperature destroy moduli spaces which can be found at zero temperature for specially tuned potentials.

This similarity is not surprising, as in some sense the theory responds in the same qualitative way to the introduction of an additional dimensionfull parameter, be it curvature or temperature. As we have mentioned however, in the case of temperature we are dealing with different states in the same theory, whereas curvature introduces a change in the definition of the theory itself. Therefore, the finite temperature analysis can tell us

something about the consistency of the theory itself, if we require that, given a consistent QFT, it should always be possible to couple it consistently to a thermal bath.

This consistency criterion puts strong constraints on theories where the vacuum state is a bouncing RG-flow. These solutions were shown to be regular and free of instabilities [12]. At finite temperature however, the phase diagram shows a jump in the free energy by a finite amount (figure 3), which should not be allowed in a consistent Maxwell construction. The interpretation is open to discussion: one possibility is that this kind of potentials lie in a holographic “swampland” and result from an inconsistent truncation of a more complete theory, e.g. a multi-field model in which the dynamics of the extra scalar cannot be neglected. Another possibility is that these models may be consistent but there are other phases, which we have neglected, and which make the phase diagram well-behaved. Since we have exhausted all spatially homogeneous solutions, one option is that these new phases break rotational or translational invariance (e.g. they may be striped phases). An instability to one such phase may be signalled by unstable quasi-normal modes, as was indeed found in models with bounces [17]. We leave these questions for future work.

We have also observed the opposite phenomenon: certain deformations of a CFT, which are not allowed at zero-temperature (because they do not lead to regular solutions), lead instead to consistent solutions above a certain T . This *thermal desingularisation* occurs in our examples around asymmetric extrema, where only one sign of the source leads to a consistent RG flow.

The possibility that solutions with “bad” singularities, in the sense of [19], may be regularised by introducing a sufficiently high² temperature was already contemplated by Gubser in that work and it would be interesting to better understand the details of this mechanism from the field theory point of view, and/or in a top-down model.

One possible field theory interpretation is that, the “wrong sign” deformation is inconsistent because of some infrared instability, which is eliminated at a sufficiently large temperature. There are perturbative QFT theory examples of such behavior, e.g. the “wrong sign” $\lambda\phi^4$ theory: at zero temperature, the would-be $\phi = 0$ “vacuum” is perturbatively unstable, but it becomes perturbatively stable at sufficiently high temperature, as we discuss in more detail in subsection 3.2

Another possibility is that the theory we study is a truncation of a more complete theory and a regular zero-temperature vacuum would be found by turning on the vevs of additional fields. If this is the case, the thermal desingularisation we found may resemble what happens with the Klebanov-Tseytlin solution [20], dual to a cascading $SU(N) \times SU(N+M)$ gauge theory: the vacuum solution is singular, but at high enough temperatures black hole solutions with a regular horizon cloaking the singularity were shown to exist [21]. However the Klebanov-Tseytlin setup does allow for a non-singular zero-temperature solution, i.e. the Klebanov-Strassler geometry, in which the IR is deformed by gaugino condensation [22], and which is the true vacuum for the theory. In our case, it would be the extra fields which are not included in the bottom-up description which would get a vev and drive the solution to a regular IR.

²As opposed to infinitesimal, as it occurs for good singularities.

A particularly interesting class of special solutions, which can be seen as separating various different branches of black holes, are the vev-driven black holes with $j = 0$. These purely vev-driven black holes are special solutions with a fixed value of the horizon parameter ϕ_h . They exist at any non-zero temperature and exhibit conformal thermodynamics. This situation, in the case of a relevant operator, bears some similarity with widely studied the case of the holographic superconductor [23, 24]. A crucial difference is that in the latter case, although no scalar couplings are turned on, in the UV there is a non-trivial UV source in the form of a chemical potential, which forces the solution to have a non-zero charge density.

Interestingly, at zero temperature, purely vev-driven holographic RG-flows are generically singular bulk solutions. In other words, for $j = 0$, the only regular solution is AdS space, with constant scalar field and $\langle O \rangle = 0$, except if $V(\phi)$ has some tuned parameters. In contrast, for $T \neq 0$, existence of regular vev-driven black holes does not require tuning the potential.

Since they have $j = 0$, vev-driven black holes satisfy the same UV boundary conditions as the AdS-Schwarzschild black holes with constant scalar field (fixed at an extremum of the potential) and the same temperature: these solutions are therefore in thermodynamic competition with each other. In all cases we have considered, it is the Schwarzschild black hole which dominates the canonical ensemble at all $T \neq 0$. It is an open (and interesting) question whether this is generic in Einstein-scalar theories, or whether there may be cases in which the vev-driven black hole is the dominant solution. Because of the relation $\langle O \rangle \propto T^\Delta$ in these solutions, these would provide a holographic example of temperature-driven condensation of a scalar operator. Examples of this kind for neutral black holes were first discussed in [25]. In the case of charged black holes they are the basis for holographic superconductors, where condensation of the scalar operator can be understood as due to an IR instability of the AdS-Schwarzschild solution [23, 24].

For those models where a regular vacuum vev-driven flow *does* exist (as in the non-generic potentials considered in [12], which allow for minimum-to-minimum holographic RG flow solutions), the situation is quite different from the one described above. At $T = 0$, there is a one-parameter family of solutions, parametrised by the arbitrary value of $\langle O \rangle$, all flowing between the same two minima of the potential. All these solutions are degenerate in free energy, therefore forming a moduli space and admit a massless dilaton excitation (the Goldstone mode of spontaneously broken conformal invariance). Going to $T \neq 0$, the entire moduli space disappears: in the example we have studied, there are no black-hole solutions with a non-trivial flow of the scalar field, which reach the same UV as the vacuum solutions. This means that finite temperature destroys the moduli space, and leaves the AdS-Schwarzschild black hole, with constant ϕ and $\langle O \rangle = 0$, as the only solution. This has an analogy in weakly-coupled field theories with a moduli space, where at finite temperature an effective potential of the form $m^2(T)\varphi^2$ may be generated, with φ a scalar representing the appropriate scalar operator. This leaves as the only minimum the one at $\langle \varphi \rangle = 0$. In holography this is also known to occur: for example, the Coulomb branch of $N = 4$ SYM has a moduli space corresponding to the fact that one can place a static extremal D3 brane anywhere within the $AdS_5 \times S^5$ geometry. This however is not possible when we go to finite

temperature and put a black hole in AdS. In all these cases the theory at finite temperature sits at the origin of the moduli space. It is an open question whether there could be models with another branch of regular vev-driven solutions, with conformal thermodynamics and $\langle O \rangle \propto T^\Delta$, which would still lift the moduli space but have a non-trivial flow.

The existence of regular $j = 0$ black holes seems tied to the presence of bounces, because the transition across $j = 0$ occurs between bouncing and non-bouncing solutions. It is unclear at present under which conditions, given a generic extremum of V , regular vev-driven black holes will or will not exist. For example, in the model with the potential shown in figure 1, purely vev-driven black holes exist which asymptote the points UV_2 and IR_1 , but not UV_1 . It would be interesting to understand what features of the scalar potential determine the existence or non-existence of these solutions.

More generally, it is an interesting but highly non-trivial question to understand which features of the potential determine whether the vacuum solution will bounce, or skip a fixed point. Although some qualitative criteria can be roughly guessed by the experience with different cases (for example, a “steeper” potential is more likely to admit bouncing vacuum flows) it would be very interesting to obtain some quantitative criteria similar to those existing for other phenomena (e.g. confinement).

Note added. During completion of this work, we became aware that a study very similar to ours was being performed independently by Y. Bea and D. Mateos [26]. The results of that work are in agreement with those presented here.

2 Holographic RG flows at finite temperature

In this section we consider the finite-temperature generalisation of the exotic RG flows found in [12]. In that paper, solutions of $(d + 1)$ -dimensional Einstein-Scalar gravity were considered, which corresponded to holographic RG flows of the dual field theory *in the vacuum*, i.e. those solution had full d -dimensional Poincaré invariance. In the following subsections we review the finite-temperature generalisations of such solutions, which contain a black hole in the interior and are only symmetric under the $d - 1$ Euclidean group plus time translations and we discuss the corresponding thermodynamics in terms of the free energy and of the thermal effective potential.

2.1 Black-hole solutions

We consider the two-derivative action of gravity coupled to a single scalar field, with a generic potential $V(\phi)$.

$$S = -M_P^{d-1} \int_{\mathcal{M}} d^{d+1}x \sqrt{-g} \left[R - \frac{1}{2} g^{ab} \partial_a \phi \partial_b \phi - V(\phi) \right] + S_{\text{GHY}} + S_{\text{ct}}, \quad (2.1a)$$

$$S_{\text{GHY}} = -2M_P^{d-1} \int_{\partial\mathcal{M}} d^d x \sqrt{\gamma} K. \quad (2.1b)$$

The sign of the action is the one appropriate for Euclidean signature. S_{GHY} is the Gibbons-Hawking-York boundary term. S_{ct} is an extra boundary term which is needed for holographic renormalisation and will be specified later.

We write the Euclidean black-hole solutions in the form:

$$ds^2 = \frac{du^2}{f(u)} + e^{2A(u)} (f(u)dt^2 + dx^i dx^i), \quad \phi = \phi(u), \quad i = 1 \dots d-1, \quad (2.2)$$

where $f(u)$ is a monotonically decreasing function taking values between zero and one, and $t \sim t + \beta$ where β is the inverse temperature.

Einstein's equations read

$$f\ddot{\phi} + (\dot{f} + d f \dot{A}) \dot{\phi} - \frac{dV}{d\phi} = 0 \quad (2.3a)$$

$$(d-1)\dot{f}\dot{A} + \left(d(d-1)\dot{A}^2 - \frac{1}{2}\dot{\phi}^2 \right) f + V(\phi) = 0 \quad (2.3b)$$

$$\ddot{f} + d\dot{A}\dot{f} = 0 \quad (2.3c)$$

$$2(d-1)\ddot{A} + \dot{\phi}^2 = 0 \quad (2.3d)$$

The equations of motion are invariant under the following transformations:

$$A \rightarrow \tilde{A} = A - \bar{A} \quad (2.4a)$$

$$u \rightarrow \tilde{u} = u + v \quad (2.4b)$$

$$(u, f(u)) \rightarrow (\tilde{u}, \tilde{f}(\tilde{u})) = (\lambda u, \lambda^2 f(u)) \quad (2.4c)$$

where \bar{A} , v and λ are constants.

The metric (2.2) describes planar black holes, whose horizon is located at u_h , i.e. $f(u_h) = 0$. Temperature and entropy density (per unit $d-1$ -volume V_{d-1}) are given, respectively, by:

$$T = \frac{e^{A(u_h)}}{4\pi} |\dot{f}(u_h)|, \quad s = 4\pi M_p^{d-1} e^{(d-1)A(u_h)}. \quad (2.5)$$

Euclidean time is compactified on a circle of length $\beta = 1/T$. The zero-temperature case (no black hole) corresponds to taking $f(u) = 1$. The AdS black-hole solution corresponds to taking constant $\phi = \phi_0$ such that $V'(\phi_0) = 0$. Defining $\ell \equiv \sqrt{-d(d-1)/V(\phi_0)}$, this solution is

$$A(u) = -\frac{u}{\ell}, \quad f(u) = 1 - e^{d(u-u_h)/\ell}, \quad \phi(u) = \phi_0, \quad (2.6)$$

and

$$T_{\text{conf}} = \frac{d}{4\pi\ell} e^{-u_h/\ell}, \quad s_{\text{conf}} = 4\pi M_p^{d-1} e^{-(d-1)u_h/\ell}. \quad (2.7)$$

In the general case, equation (2.3c) can be integrated once to obtain

$$\dot{f}(u)e^{dA(u)} = -D, \quad (2.8)$$

where D is a non-negative integration constant. We can relate it to the black-hole temperature and entropy by evaluating equation (2.8) at the horizon $u = u_h$ and using equation (2.5), leading to

$$D = \frac{Ts}{M_p^{d-1}}. \quad (2.9)$$

We restrict to solutions which reach an asymptotically AdS region (UV) where the scalar field approaches a maximum ϕ_{UV} (which without loss of generality we take to be at $\phi = 0$) of the potential, i.e.

$$V(\phi) \simeq -\frac{d(d-1)}{\ell^2} + \frac{m^2}{2}\phi^2 + \dots \tag{2.10}$$

In this region the solution takes the asymptotic form as $u \rightarrow -\infty$,

$$\phi = \phi_- e^{\Delta_- u/\ell} + \dots + \phi_+ e^{\Delta_+ u/\ell} + \dots, \quad \Delta_{\pm} = \frac{d}{2} \pm \frac{1}{2}\sqrt{d^2 + 4m^2\ell^2}, \tag{2.11a}$$

$$A(u) = -\frac{u}{\ell} + \dots \tag{2.11b}$$

$$f(u) = 1 - \frac{\ell D}{d} e^{du/\ell} + \dots \tag{2.11c}$$

The parameters ϕ_- and ϕ_+ are related to the UV coupling j and to the vev of the dual operator O (whose dimension³ is $\Delta = \Delta_+$) by

$$\phi_- = j \ell^{\Delta_-}, \quad \phi_+ = \frac{\langle O \rangle \ell^{\Delta_+}}{(M_p \ell)^{d-1} (d - 2\Delta_-)}. \tag{2.12}$$

Finally, the constant D is related to the temperature and entropy by equation (2.9). Generically $\phi_- \neq 0$ and the flows are driven by a deformation of the CFT by adding a source to a relevant operator. For special solutions with $\phi_- = 0$ the asymptotic expansion starts at sub-leading order with ϕ_+ . These flows are driven by a vev of the dual operator, and as we will see they play a special limiting role in the space of solutions.

2.2 Dimensionless thermodynamic parameters

It is useful to classify black-hole solutions in terms of a dimensionless and diff-invariant quantity. One useful choice is the horizon value of the scalar field,

$$\phi_h \equiv \phi(u_h). \tag{2.13}$$

As was shown in [17, 27] and explained in detail in appendix A, ϕ_h determines the dimensionless quantity

$$\mathcal{T} \equiv \frac{T}{|j|^{1/\Delta_-}}. \tag{2.14}$$

Thus, for each ϕ_h there is a one-parameter family of black-hole solutions with fixed \mathcal{T} , and fixing either the temperature or the UV source j selects a single solution in this family. In other words, we can build a map

$$(\phi_h, j) \rightarrow T(\phi_h, j) = j^{\frac{1}{\Delta_-}} \mathcal{T}(\phi_h). \tag{2.15}$$

³We will only discuss “standard quantisation”, where the dimension of the deforming operator $\Delta \geq d/2$.

In what follows, we will use \mathcal{T} itself, rather than ϕ_h , as an independent parameter,⁴ since the former is directly related to the boundary quantities T and j . Similarly, any physical quantity measured in units of j (entropy, vev, etc) only depends on \mathcal{T} (or equivalently on ϕ_h). This is the case in particular for the rescaled entropy density

$$\mathcal{S}(\mathcal{T}) \equiv \frac{s}{|j|^{(d-1)/\Delta_-}}. \quad (2.16)$$

2.3 First order formalism

To classify black-hole solutions we will often resort to a first order formalism. This is a finite-temperature extension, developed in detail in appendices A and C, of the standard first order formulation of holographic RG flows, [8, 28]. The formulation we use here is different from the first order formalism one obtains from Hamilton-Jacobi theory, in which the function A , f and ϕ all satisfy first order gradient flow equations [29]. In particular, the superpotential $W(\phi)$ is not a solution of the Hamilton-Jacobi equation, which at finite temperature is not separable.

Black-hole solutions can be classified in terms of a *superpotential*, a function $W(\phi)$ which determines the scale factor and scalar field profile by the equations

$$\dot{A}(u) = -\frac{1}{2(d-1)}W(\phi(u)), \quad \dot{\phi}(u) = \frac{dW}{d\phi}(\phi(u)). \quad (2.17)$$

The superpotential $W(\phi)$ and the blackness function $f(u)$ (more precisely, the function $f(\phi)$ defined such that $f(u) = f(\phi(u))$) are determined together by solving a coupled non-linear third-order system of equations in the independent variable ϕ , equations (A.2)–(A.3). Because this system couples W to f , $W(\phi)$ depends on temperature.

Below we list the most important properties of the superpotential $W(\phi)$ which we will be useful for our analysis

1. Imposing regularity at the horizon, the functions $W(\phi)$ and $f(\phi)$ are completely specified by assigning a single parameter, i.e. the value ϕ_h of the scalar field at the horizon, or equivalently (at least piecewise) the dimensionless temperature parameter \mathcal{T} .
2. The superpotential is monotonically increasing along the flow, as a function of the holographic coordinate u . However it can be multi-valued as a function of ϕ if the scalar field profile $\phi(u)$ is non-monotonic [12].
3. For source-driven flows, i.e. solutions with $j \neq 0$ in equation (2.11a), the superpotential (denoted in this case by $W^-(\phi)$) takes the form of a universal (i.e. \mathcal{T} -independent) analytic expansion around the boundary value $\phi = 0$, plus a one-parameter, \mathcal{T} -dependent, non-analytic contribution controlled by an integration constant $C(\mathcal{T})$ (or equivalently, $C(\phi_h)$),

$$W^-(\phi) = \frac{2(d-1)}{\ell} + \frac{\Delta_-}{2\ell}\phi^2 + \mathcal{O}(\phi^4) + \frac{C(\mathcal{T})}{\ell}|\phi|^{d/\Delta_-} \left[1 + \mathcal{O}(\phi)\right] + \dots \quad (2.18)$$

⁴Notice that the map $\mathcal{T}(\phi_h)$ is singled-valued, but not necessarily invertible. Therefore, we can use \mathcal{T} as an independent parameter piecewise, i.e. there may be more than one black-hole branch with the same values of \mathcal{T} but different ϕ_h . This is usually the case when using temperature to parametrise asymptotically AdS black holes.

where the ellipsis refers to terms of higher non-analytic order [31, 32] but with no new free parameter. The value of the source j enters the full solution as the integration constant of the flow equation (2.17) for ϕ . The quantity $C(\mathcal{T})$ determines the sub-leading term ϕ_+ in the scalar field expansion, and consequently the vev of the dual operator by equation (2.12), by

$$\phi_+ = \frac{d}{\Delta_-} \frac{C(\mathcal{T})}{(d - 2\Delta_-)} |j|^{\Delta_+/\Delta_-} \text{sign}(j). \quad (2.19)$$

4. For vev-driven flows ($j = 0$) the superpotential (denoted in this case by $W^+(\phi)$) consists of a purely analytic expansion in ϕ , with no additional \mathcal{T} -dependent deformation parameters,

$$W^+(\phi) = \frac{2(d-1)}{\ell} + \frac{\Delta_+}{2\ell} \phi^2 + \mathcal{O}(\phi^4) \quad (2.20)$$

In this case the integration constant of the first order equation for $\phi(u)$ is ϕ_+ , which is a free parameter for these solutions.

Alternatively, as explained in detail in appendix C one can define the *scalar variables* (that transform as scalars under a diffeomorphism of u):

$$X(\phi) \equiv \frac{1}{\sqrt{2d(d-1)}} \frac{\dot{\phi}}{\dot{A}(u)}, \quad Y(\phi) \equiv \frac{1}{d} \frac{\dot{g}(u)}{\dot{A}(u)}, \quad (2.21)$$

where the function g is defined as $g = \log f$. Then Einstein's equations can be reduced to two coupled first order equations for X and Y as detailed in appendix C. The functions X and Y contain all physically relevant information on the system both in the vacuum and at finite temperature. For example the free energy can be read off directly from the boundary asymptotics of the functions X and Y . One can think of the boundary values of Y and X as the enthalpy s/T and a combination of energy with the enthalpy, respectively. The first order formalisms in terms of the superpotential and the scalar variables are completely equivalent, e.g. X is the logarithmic derivative of W .

2.4 The free energy

The free energy associated to the solution is given by the Euclidean renormalised on-shell action,

$$\beta\mathcal{F} = S_{\text{on-shell}}^{(\text{ren})} \quad (2.22)$$

Here we will focus on source-driven flows leaving the special case of vev-driven flows for a later section. An explicit calculation, which is performed with two independent methods in appendices B and C.3, leads to the expression

$$\frac{\mathcal{F}}{V_{d-1}} = -\frac{Ts}{d} - (M_p \ell)^{(d-1)} C(\mathcal{T}) |j|^{d/\Delta_-}, \quad (2.23)$$

where T is the black-hole temperature, s the BH entropy density and $C(\mathcal{T})$ is the parameter controlling the sub-leading asymptotics of the superpotential near the boundary, see equation (2.18).

In the canonical ensemble we are using, the boundary data which are kept fixed are T and j , and $\mathcal{F} = \mathcal{F}(T, j)$. Black-hole thermodynamics implies that the black-hole entropy density s and energy density ϵ are

$$s = -\frac{1}{V_{d-1}} \frac{\partial \mathcal{F}}{\partial T}, \quad \epsilon = \frac{\mathcal{F}}{V_{d-1}} + Ts \quad (2.24)$$

The dual operator vev is the conjugate variable to j ,

$$\langle O \rangle(T, j) = -\frac{1}{V_{d-1}} \frac{\partial \mathcal{F}(T, j)}{\partial j} \quad (2.25)$$

and it can be shown (see appendix B.2) that the right-hand side of the equation above agrees exactly with the definition of $\langle O \rangle$ from the holographic dictionary, equation (2.19). The pressure $p = \mathcal{F}/V_{d-1}$ obeys the differential relation

$$dp = -sdT - \langle O \rangle dj. \quad (2.26)$$

Taking the conformal limit $j \rightarrow 0$ in (2.23) we recover conformal thermodynamics, $\mathcal{F} = V_{d-1}Ts/d$, which implies $s \propto T^{d-1}$ as found by integrating the differential relation (2.24). Finally, at $T = 0$ we recover the known result [31, 33]:

$$\mathcal{F}_{T=0} = -(M_p \ell)^{d-1} V_{d-1} C_0 |j|^{d/\Delta_-} = -\frac{\Delta_-}{d} V_{d-1} \langle O \rangle_{T=0} j, \quad (2.27)$$

where C_0 is the value of $C(\mathcal{T})$ in the zero-temperature vacuum.

It is convenient to rewrite the free energy in terms of the dimensionless variables introduced in section 2.2,

$$\frac{\mathcal{F}(T, j)}{V_{d-1}} = T^d \left(-\frac{\sigma(\mathcal{T})}{d} + \gamma(\mathcal{T}) \right), \quad (2.28)$$

where we have defined the dimensionless quantities

$$\sigma(\mathcal{T}) \equiv \frac{s}{T^{d-1}} = \frac{\mathcal{S}(\mathcal{T})}{\mathcal{T}^{(d-1)}}, \quad \gamma(\mathcal{T}) \equiv (M_p \ell)^{d-1} \frac{C(\mathcal{T})}{\mathcal{T}^d}. \quad (2.29)$$

Apart from the overall T^d scaling, all the non-trivial dependence on T and j in the free energy only appears through the combination \mathcal{T} defined in equation (2.14).

2.5 The thermal effective potential

From the free energy, we can define the finite temperature effective potential by a Legendre transform. First, we trade j for its conjugate variable, i.e. the dual operator vev O (in this section we omit the brackets for simplicity of notation): inverting the relation (2.25) to obtain $j(T, O)$, the effective potential is then defined as

$$V_{\text{eff}}(T, O) = \mathcal{F} + Oj, \quad dV_{\text{eff}} = V_{d-1} \left[-sdT + jdO \right], \quad (2.30)$$

and it satisfied the relations:

$$\frac{1}{V_{d-1}} \frac{\partial V_{\text{eff}}}{\partial O} = j, \quad \frac{1}{V_{d-1}} \frac{\partial V_{\text{eff}}}{\partial T} = -s. \quad (2.31)$$

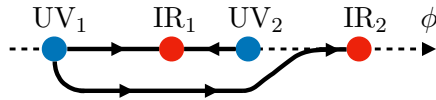


Figure 4. Schematic structure of a field theory which presents multiple RG-flows: in particular, there are two flows starting at the fixed point UV_1 , one going to the closer IR fixed point IR_1 , the second skipping IR_1 and ending at IR_2 . On the other hand there is no flow from UV_2 to IR_2 .

As for the temperature, we can introduce a dimensionless vev parameter,

$$\mathcal{O} = \frac{O}{T^{\Delta_+}}. \quad (2.32)$$

Starting from equation (2.28) it is then easy to show that one can write (2.30) in the form

$$V_{\text{eff}}(T, \mathcal{O}) = T^d \mathcal{V}(\mathcal{O}), \quad (2.33)$$

where $\mathcal{V}(\mathcal{O})$ is the Legendre transform of \mathcal{F}/T^d with respect to $\mathcal{T}^{-\Delta_-}$ (which is indeed the dual variable to \mathcal{O}).

Equation (2.33) is useful because it allows the treatment of theories in which the source $j = 0$, i.e. the case of pure vev flows: these are the values of \mathcal{O} which extremise the function \mathcal{V} . We will see examples of these flows in the following sections, and we will discuss them in detail in section 5.

3 Thermal phase transitions in multi-vacuum theories

In the previous section we have developed a general expression for the free energy of any black-hole solution, in terms of the UV source, temperature, entropy and vev of the operator dual to ϕ .

We are now ready to study the phase diagram of black-hole solutions in situations where the zero-temperature RG flow displays exotic features.

In this section we concentrate on situations where *skipping flows* are present: in the presence of several maxima and minima of the scalar potential, these are flows which skip an intermediate potential IR fixed point and end at a fixed point further away in field space, as schematically represented in figure 4.

Vacua of the dual field theory correspond to IR-regular flows. In the models at hand there may be multiple distinct IR-regular RG flows with the same UV boundary conditions. These are interpreted, under the holographic map, as different vacua of the same theory, with different β -functions and different IR endpoints. At zero temperature, the true vacuum is the one with the lowest free energy. In [12] it was shown that this is the flow where the parameter $C(T = 0)$ is the largest, i.e. the one with the largest vev at fixed source j (cfr. equation (2.27)). This guarantees that the relevant solution has the lowest free energy.

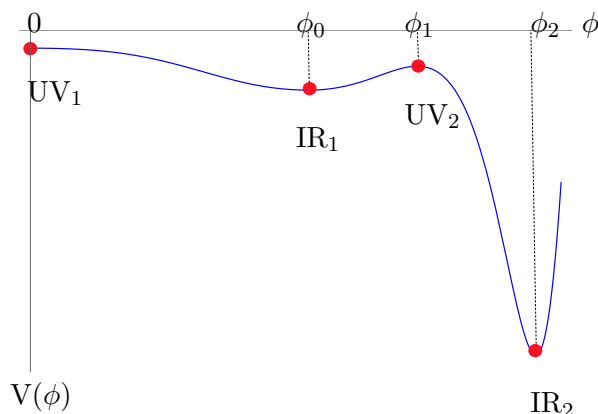


Figure 5. Plot of the degree-12 potential which allows skipping solutions. This potential has several extrema which are denoted as UV_1 , UV_2 , IR_1 and IR_2 .

3.1 Skipping RG flows at zero-temperature

As an example of the behaviour described above, in [12] the following 12th-order potential was considered,

$$V(\phi) = -\frac{d(d-1)}{\ell^2} + \int_0^\phi V'(x)dx. \quad (3.1)$$

where

$$\ell^2 V'(\phi) := -\phi(\phi^2 - \phi_0^2)(\phi^2 - \phi_1^2)(\phi^2 - \phi_2^2)(\phi^2 - \phi_3^2) \left(\phi^2 - \frac{\Delta(\Delta-d)}{\phi_0^2 \phi_1^2 \phi_2^2 \phi_3^2} \right), \quad (3.2)$$

The potential has extrema at the points $0 < \phi_0 < \phi_1 < \phi_2 < \phi_3$. We make the specific choice:

$$\begin{aligned} d = 4 & \quad \phi_0 = 1.0837 & \phi_1 = 1.1316 \\ \Delta = 2.8 & \quad \phi_2 = 1.9200 & \phi_3 = 2.1500. \end{aligned} \quad (3.3)$$

With these choices, the operator dimensions at the various fixed points are given by:

$$\Delta_{UV_1} = 2.8, \quad \Delta_{UV_2} = 3.1, \quad \Delta_{IR_1} = 4.5, \quad \Delta_{IR_2} = 11.6. \quad (3.4)$$

The potential is shown in figure 5, where the correspondence between the values ϕ_i and the UV and IR fixed points is made manifest. The explicit expression of $V(\phi)$ can be found in appendix A of [12].

At $T = 0$ there are several IR-regular RG-flow solutions, displayed schematically in figure 6, which shows the zero-temperature superpotential of each flow.⁵

- $UV_1 \rightarrow IR_1$. This solution is the standard holographic RG flow which connects a maximum of the potential (UV) to the nearest minimum (IR).

⁵Figure 6 displays the *superpotentials* and the critical curve $B \sim \sqrt{-V}$ (which bounds the space of solutions, see [12]). Therefore what is presented as a maximum (minimum) in that figure actually corresponds to a minimum (maximum) of the potential in figure 5.

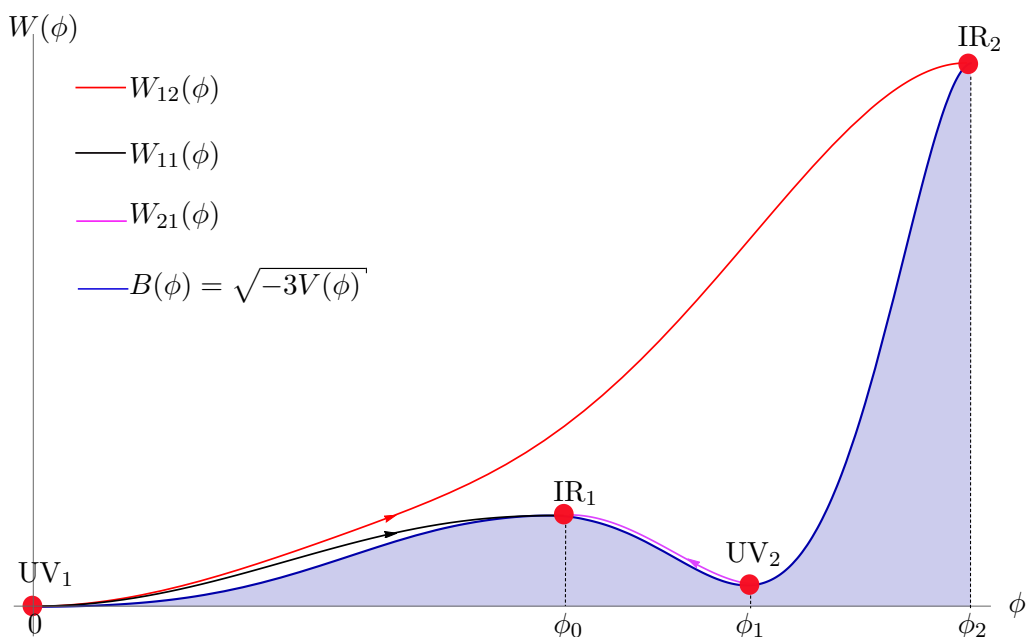


Figure 6. The vacuum RG flow solutions arising from the potential (3.1). All flows interpolate between one of the UV fixed points and one of the IR fixed points. The plotted lines are the superpotentials corresponding to each flow (see appendix A). The arrows represent the direction of the flow from the UV to the IR. The blue area is the forbidden region below the curve $B(\phi) = \sqrt{-4(d-1)V(\phi)/d}$ (where we set $d = 4$), which bounds from below any solution to the superpotential equation at $T = 0$ (as explained in appendix A.2).

- $UV_1 \rightarrow IR_2$. This solution on the other hand skips the first minimum and ends at the next available IR fixed point at IR_2 . This kind of solution is not found in *generic* potentials admitting several extrema: for it to exist the extremum IR_1 has to be sufficiently shallow.
- $UV_2 \rightarrow IR_1$. This solution corresponds to a standard flow with a *negative* source from the second maximum of $V(\phi)$, reaching the closest available IR fixed point. Notice that there is no solution connecting UV_2 to IR_2 . The reason is that there already is a regular flow arriving at IR_2 (the one from UV_1 : since flows reaching (from a given direction) a minimum of the potential are isolated, this prevents other flows to reach the same IR).

The two solutions leaving UV_1 , correspond to two vacua of the same UV theory. The one with the lowest free energy (2.27) is the skipping one $UV_1 \rightarrow IR_2$, since it has the largest vev parameter C_0 , as can be seen immediately from the fact that the corresponding superpotential (W_{12} in figure 6) increases faster close to the origin.

3.2 Finite temperature solutions

We now move to finite T by considering black-hole solutions, of the form (2.2) in the model with the same potential in figure 5. These black holes are uniquely characterised by a

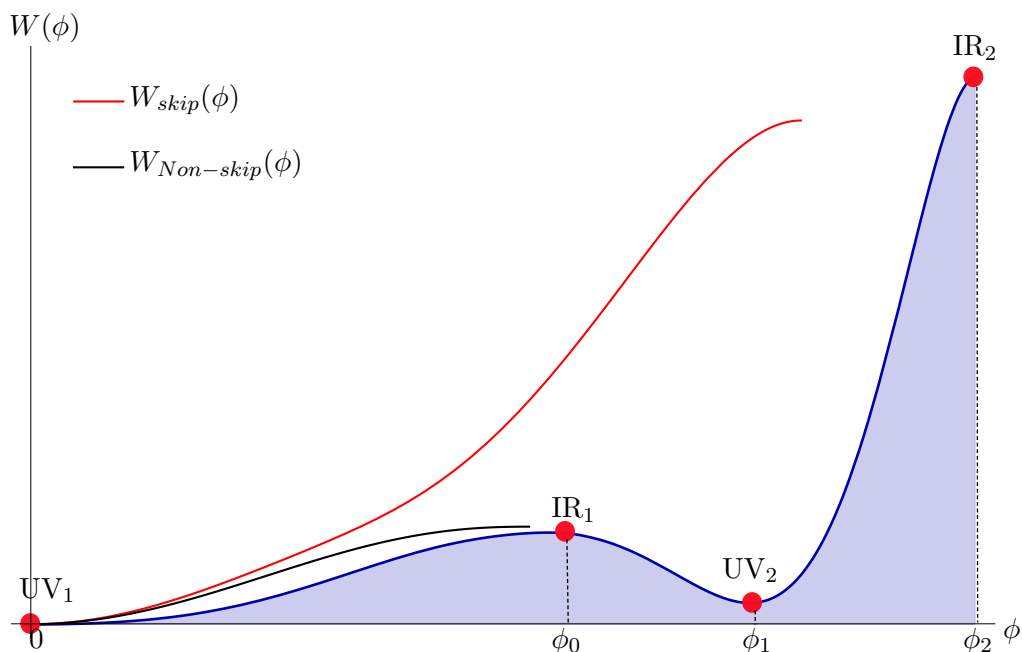


Figure 7. Finite temperature solutions arising from the potential (3.1) and connecting to UV_1 .

dimensionless number: the value ϕ_h of the scalar field at the horizon. If we keep the value of the UV source j fixed, we expect ϕ_h to determine all other quantities, (temperature, entropy, free energy, etc.) Therefore we are interested in constructing all solutions with ϕ_h ranging from zero (UV_1) to ϕ_2 (IR_2). Indeed, we will see that for every value of ϕ_h there exists at most one black-hole solution with all other UV data fixed.

It has to be stressed that, in order to be considered as different states in the same dual QFT, two solutions must connect to the same UV fixed point, with the same value of the source parameter j .

As we will see below, depending on the value reached at the horizon by the scalar field, integrating the solution “backwards” away from the horizon may lead either to UV_1 or to UV_2 . These represent two disconnected classes of solutions, since they have different boundary conditions at the UV boundary. From the dual field theory standpoint, they represent thermal states in different (deformed) CFTs. For this reason, we analyse them separately.

Flows from UV_1

We first consider solutions connecting to the UV_1 fixed point at $\phi = 0$. The two zero-temperature vacuum flows are the black and red curves in figure 6. Turning on temperature, the situation is represented in figure 7, where the endpoint of each flow is now at the horizon, where $\phi = \phi_h$.

There are now up to *three* branches of solutions at fixed \mathcal{T} , whose (dimensionless) temperature and entropy density as a function of the horizon value ϕ_h are is represented in figure 8. The corresponding vev parameter $C(\phi_h)$ is shown in figure 9. As one can

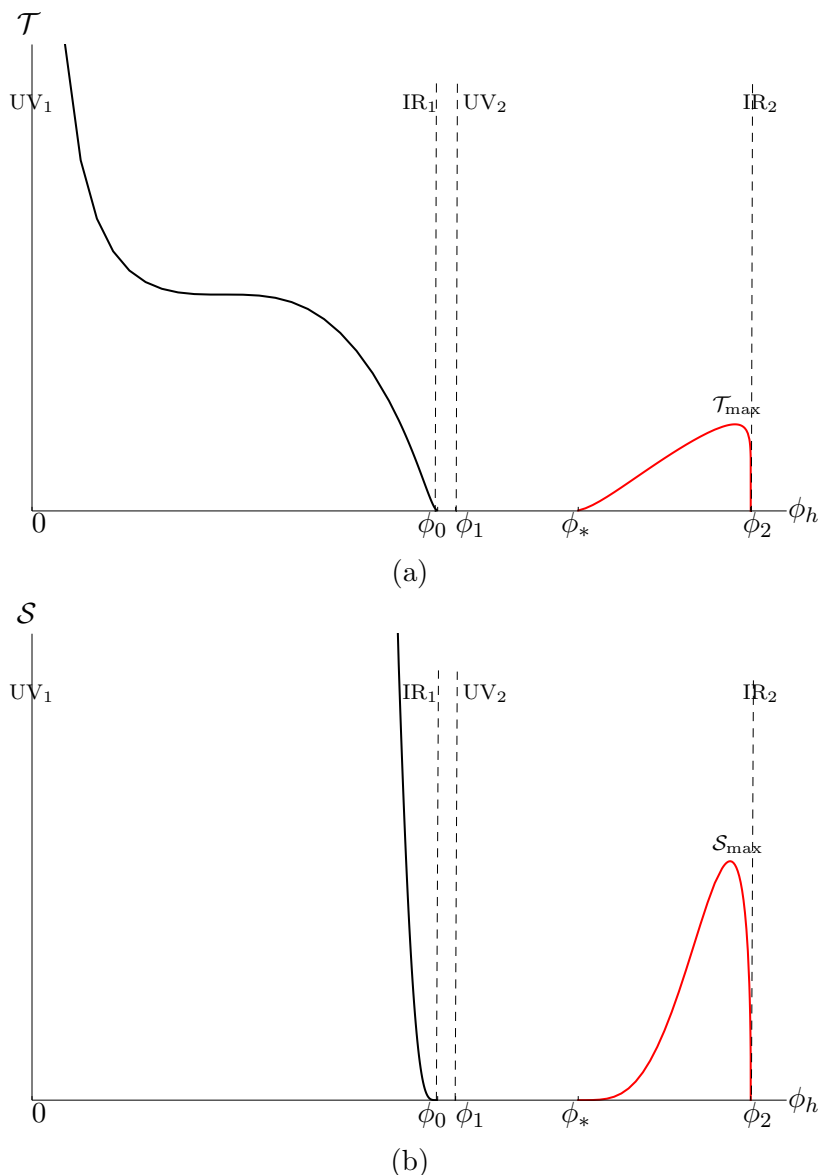


Figure 8. The dimensionless temperature (a) and the entropy density (b) as a function of the scalar field horizon value ϕ_h for black holes connecting to UV_1 .

observe, there is a range of horizon values, ϕ_h , (between ϕ_0 and a critical point which we denote by ϕ_*) for which no solution exists which continuously connects to UV_1 .

1. *Solutions with $\phi_h < \phi_0$.* These are black holes continuously connected to the non-skipping vacuum flow from $UV_1 \rightarrow IR_1$. As the temperature is increased, the horizon moves closer and closer to the UV fixed point of the vacuum solution. This is the standard behaviour at finite temperature for the simplest RG flows, connecting two consecutive extrema of the scalar potential.
2. *Solutions with $\phi_* < \phi_h < \phi_2$.* These solutions all skip IR_1 and flow to the region between UV_2 and IR_2 . As one can see from figure 8, these solutions have a maximal

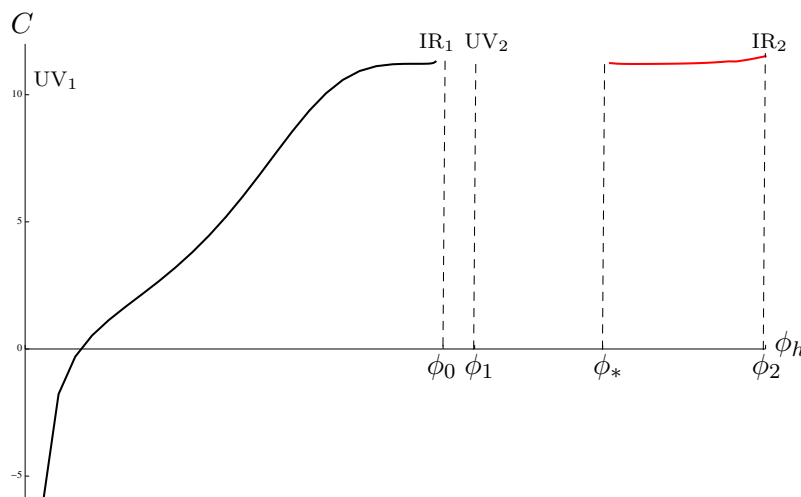


Figure 9. The (dimensionless) vev parameter C as a function of the scalar field horizon value ϕ_h for black holes connecting to UV_1 . As the horizon approaches the UV fixed point at $\phi = 0$, the parameter $C \rightarrow -\infty$.

temperature \mathcal{T}_{\max} and entropy \mathcal{S}_{\max} . For each $\mathcal{T} < \mathcal{T}_{\max}$ there are two solutions. Of the two, the one with the larger ϕ_h is the deformation of the zero-temperature skipping flow $UV_1 \rightarrow IR_2$, for which the temperature increases as the horizon moves away from the IR fixed point. The second solution is a new branch, which has no zero-temperature analogue, and for which the temperature increases as the horizon moves *towards* the IR fixed point IR_2 . At the critical value corresponding to \mathcal{T}_{\max} , the two solutions merge. Both branches extend to arbitrarily low temperature, but only one of them (the one with higher ϕ_h) actually connects to a horizonless zero- T solution. The fate of the other branch as $\mathcal{T} \rightarrow 0$, $\phi \rightarrow \phi_*^+$, will be discussed separately at the end of this section.

Given the situation described above, it is clear that at $\mathcal{T} > \mathcal{T}_{\max}$ there is a unique black hole solution, which belongs to the non-skipping branch. At zero temperature however, as we discussed at the beginning of this section, the true ground state is the skipping branch which reaches IR_2 . Therefore, a skipping solution is expected to continue to be the ground state for small but finite temperature. In other words, there should be a phase transition between the skipping and non-skipping black holes at some finite temperature $\mathcal{T}_c < \mathcal{T}_{\max}$. This is indeed confirmed by a numerical analysis, and it is clearly visible in figure 10, where we display the free energy, as a function of the temperature, of the three branches of solutions connecting to UV_1 .

Flows from UV_2

Solutions with scalar field horizon values ϕ_h in the range $\phi_0 < \phi_h < \phi_*$ connect to UV_2 , rather than UV_1 . This explains the empty gap in horizon values in figures 8 and 9. These solutions belong to a different dual field theory (a deformation of a different UV CFT) from those flowing from UV_1 .

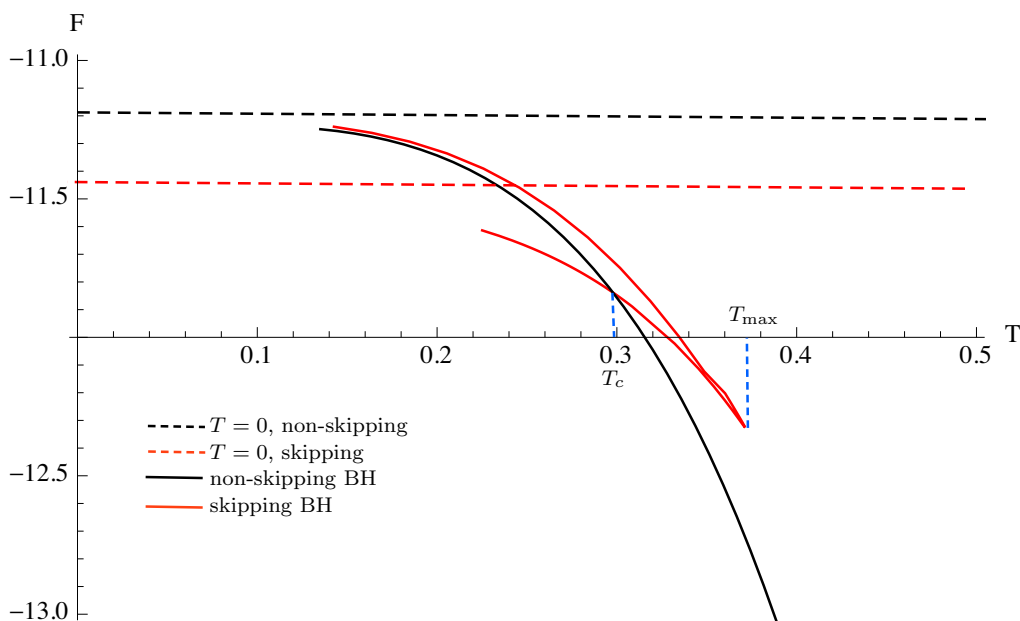


Figure 10. The figure shows the free energy density as a function of temperature for the solutions connecting to UV_1 . The units were fixed such that $j = 1$ (so that $\mathcal{T} = T$). The black curve corresponds to the non-skipping branch, and it extends from $T = 0$ to infinity. The red curves are the two skipping branches. The dashed horizontal lines represent the free energy of the zero-temperature skipping (red) and non-skipping (black) solutions. Of the two skipping black-hole solutions, the one with lower free energy connects to the true vacuum skipping flow as $T \rightarrow 0$ (The curves shown here do not reach all the way to zero temperature due to limitations in the numerics). Both skipping branches disappear at a maximal temperature, which is numerically found to be $T_{\max} \simeq 0.37$ in units where $j = 1$. The phase transition occurs at the temperature $T_c \simeq 0.3$, where the skipping and non-skipping branches cross.

The set of flows emerging from UV_2 is represented schematically in figure 11. These flows can be divided into three different classes: those with positive source (represented in blue), those with negative source which *bounce* (i.e. where the scalar field inverts its direction along the flow, dashed purple curve) and do not bounce (solid purple curve). Bouncing solutions of this kind were discussed at zero-temperature in [12], and examples were studied in a model with a different bulk potential, which will be the subject of section 4. Here we see a new feature: although in the current example with the potential in figure 5 there are no bouncing solutions at zero-temperature, these may appear at finite temperature. A similar phenomenon was observed in [16] in the case of a non-zero boundary curvature.

All these solutions can be classified according to the endpoint value ϕ_h . Their dimensionless temperature, entropy density, and vev parameter are represented as a function of ϕ_h in figures 12 and 13 (the complement of figure 8 and 9). Notice the existence of a special point ϕ_c , whose value lies between ϕ_1 and ϕ_* , which separates positive-source and negative-source black holes.

1. *Solutions with $\phi_0 < \phi_h < \phi_1$ (solid purple curve in figures 11 and 12).* These black holes are the finite temperature continuations of the vacuum flow from UV_2 to IR_1

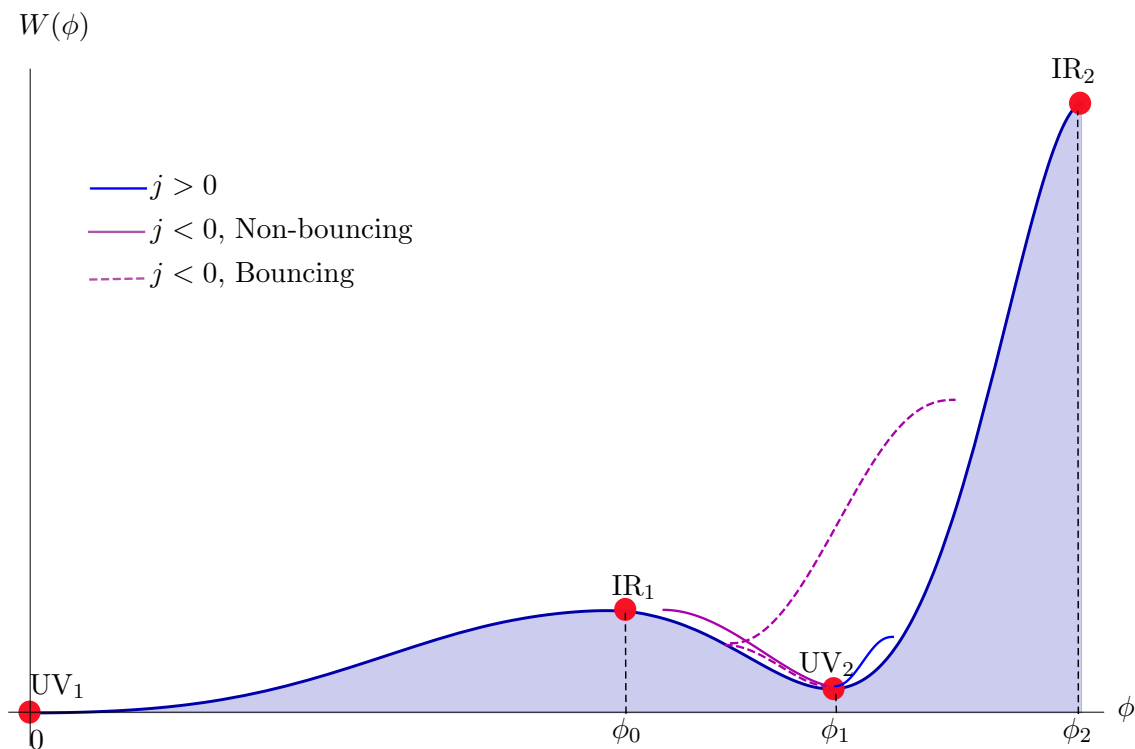


Figure 11. Finite temperature solutions arising from the potential (3.1) and connecting to UV_2 .

shown in figure 6. As the temperature is increased, the scalar field endpoint moves from IR_1 to higher values towards UV_2 . These flows have a *negative* value of the UV deformation parameter j , since $\phi(u)$ decreases along the flow.

2. *Solutions with $\phi_1 < \phi_h < \phi_c$ (blue curve in figures 11 and 12).* These solutions flow from UV_2 to larger values of ϕ , and therefore, although they originate from the same fixed point, they do not belong to the same class of deformed CFT as those in the previous class, since they differ by the sign of the source of the deformation. Solutions in this class have no zero-temperature analogue (recall that a regular flow from UV_2 with positive source does not exist at zero temperature), and in fact they only exist above a *minimum* temperature T_{\min} , as can be seen from figure 12. At both ends of this range, the temperature asymptotes to infinity.
3. *Solutions with $\phi_c < \phi_h < \phi_*$ (dashed purple curve in figures 11 and 12).* These solutions display a bounce in the flow: they start out from UV_2 with *decreasing* scalar field, (i.e. they have $j < 0$ and therefore belong to the same UV theories as those in class 1 above). Before they reach IR_1 , the scalar field reaches a minimum, the flow inverts its direction and starts running towards IR_2 . The horizon lies somewhere in between the critical points ϕ_c and ϕ_* , beyond which we find solutions in the other classes.

Below we discuss several properties of the space of solutions which asymptote UV_2 .

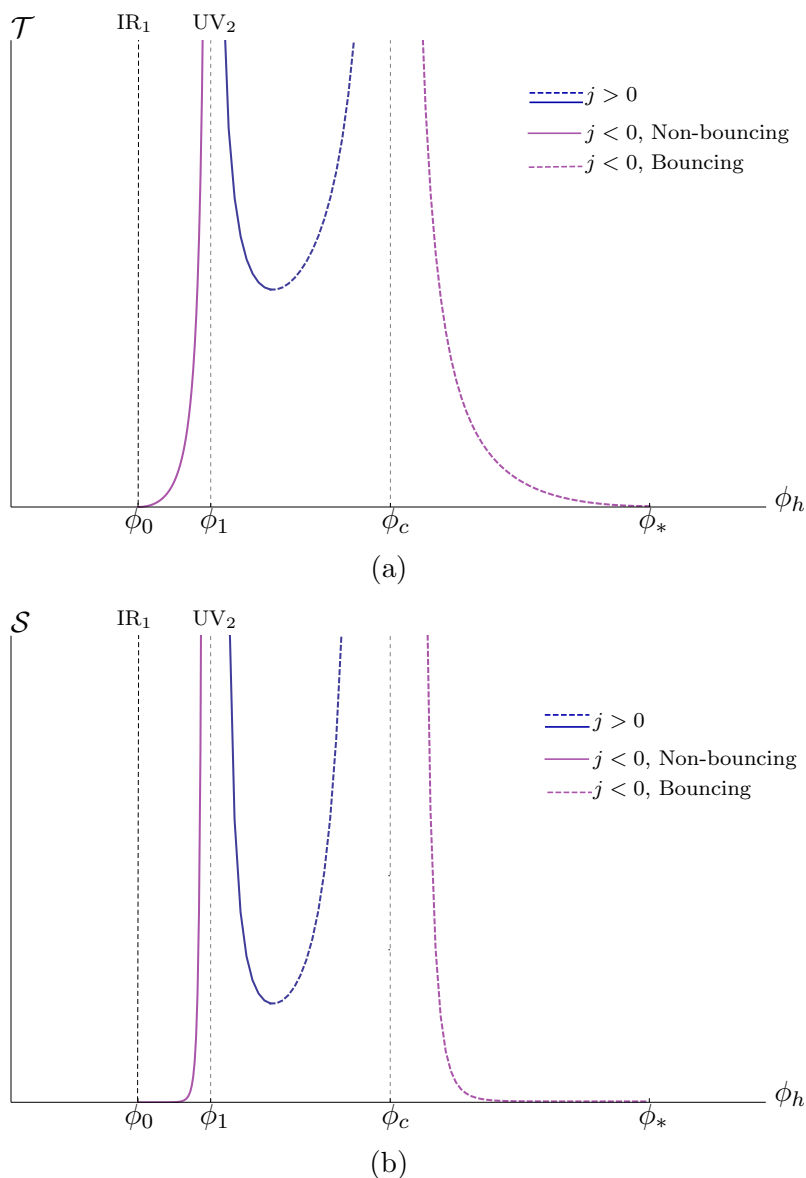


Figure 12. The temperature (a) and entropy density (b) as a function of the scalar field horizon value ϕ_h for black holes connecting to UV_2 . The colours and dashes correspond to the flows shown in figure 11.

Negative source phase diagram. From figure 12 (a) we see that, for $j < 0$, there are two black-hole solutions for any given temperature $T > 0$. By computing their free energies numerically, we have shown that the thermodynamically favoured black hole (the one with lowest free energy) is, at all temperatures, the non-bouncing solution, i.e. the one with horizon between IR_1 and UV_2 , which continuously connects to the vacuum flow ending at ϕ_0 at zero temperature. This result is shown in figure 14, where we plot the free energy difference between the non-bouncing and bouncing solutions.

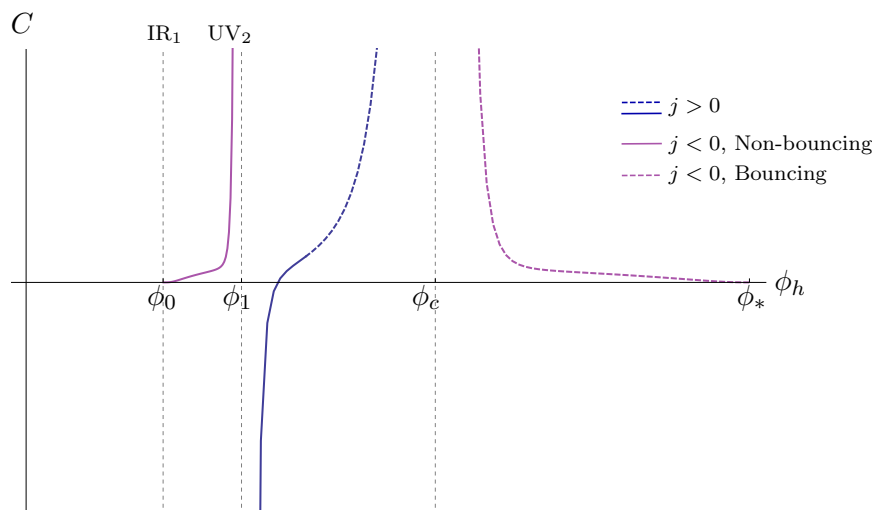


Figure 13. The dimensionless vev parameter C as a function of the scalar field horizon value ϕ_h for black holes connecting to UV_2 .

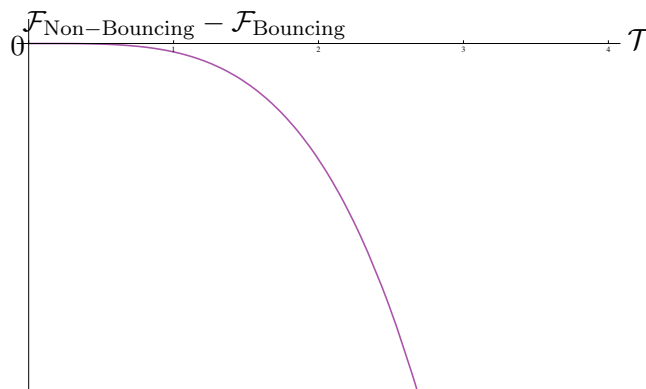


Figure 14. Free energy difference as a function of the (dimensionless) temperature between the two branches of black holes emanating from UV_2 with negative source (dashed and solid purple curves in figures 11 and 12).

Positive source phase diagram. We now turn to the black-hole solutions starting at UV_2 with positive source. Recall that no such regular solutions exist at zero temperature: as one can see from figure 6 the only regular flow vacuum flow from UV_2 is the one ending at IR_1 , and it has negative source. As was already noted in [12], this means that for $j > 0$ the dual field theory is ill defined.⁶

Interestingly, at finite temperatures larger than a minimal temperature T_{\min} , black-hole solutions with $j > 0$ start to exist, as one can see in figure 12. In other words, it is only by turning on a sufficiently high temperature that we obtain regular solutions with this sign of the source. A possible interpretation from the field theory point of view would be

⁶examples of this behaviour in perturbative field theories are common, e.g $\lambda\phi^4$ theory is defined only for $\lambda > 0$, and the same can be said for the 't Hooft parameter $\lambda = g^2 N$ in Yang-Mills theory.

the fact that temperature provides an IR regulator which eliminates some IR pathologies which made the vacuum theory ill-defined.

In field theory, there exist weakly coupled examples of this behavior. Consider a unstable scalar theory with a potential

$$V(\phi) = -\frac{\lambda}{4}\phi^2 - \frac{m^2}{2}\phi^2, \quad m^2 > 0, \quad \lambda > 0 \quad (3.5)$$

It is clear that at zero temperature, this theory does not exist as the potential is unstable. However at finite temperature there is a one-loop thermal correction to the mass, that with the appropriate field content is positive, so now the effective potential becomes

$$V_T(\phi) = -\frac{\lambda}{4}\phi^2 - \frac{m^2}{2}\phi^2 + \frac{\kappa}{2}T^2\phi^2, \quad \kappa > 0 \quad (3.6)$$

Now, for $T > \frac{m}{\sqrt{\kappa}}$, this effective potential has three extrema. One is still at $\phi = 0$, but this extremum is now stable. There are two more extrema at $\phi = \pm\sqrt{\frac{\kappa T^2 - m^2}{\lambda}}$ where ϕ has a non-trivial vev but they are unstable. The theory is perturbatively stable around $\phi = 0$ (but of course non-perturbatively unstable). Moreover, the stable minimum with zero vev has the lowest energy of all three minima.

This can be made into a large N adjoint theory using N_B adjoint matrices ϕ^I , as well as N_F adjoint fermions ψ^I , with an action of the form

$$S = \int d^4x \left[\frac{1}{2} \sum_{I=1}^{N_B} \text{Tr} [\partial\phi^I \partial\phi^{I,\dagger}] + \sum_{I=1}^{N_F} \text{Tr} [\bar{\psi}^I \not{\partial}\psi^I] + \sum_{I,J,K} (g_{IJK} \text{Tr} [\phi^I \bar{\psi}^J \psi^K] + cc) - \sum_{I=1}^{N_B} \left[-\frac{\lambda}{4} \text{Tr} [(\phi^I \phi^{I,\dagger})^2] + \frac{\kappa T^2 - m^2}{2} \text{Tr} [\phi^I \phi^{I,\dagger}] \right] \right] \quad (3.7)$$

where the numbers $N_{B,F}$ can be adjusted so that the one-loop correction $\kappa > 0$.

Returning to our phase diagram, we observe that at any temperature above T_{\min} there are two black-hole solutions. Computing their free energy, we found that the one which dominates the ensemble is always the one with smaller ϕ_h (the solid blue branch in figure 12). This is shown in figure 15. The dominant branch is the one which, as the temperature rises, approaches the AdS-Schwarzschild black hole with constant scalar field $\phi = \phi_1$ (see discussion in the following paragraphs).

Infinite temperature limits. There are three situations in which the dimensionless temperature $\mathcal{T} \rightarrow \infty$: when the horizon approaches one or the other UV fixed points, and at the point ϕ_c . The first case is easy to understand, as it corresponds to the limit of large AdS-Schwarzschild black holes in which the scalar is approximately constant. The divergence of the temperature at $\phi_h = \phi_c$ is more interesting: across this point the UV source changes sign. From the definition (2.14), the divergence in \mathcal{T} can be interpreted in this case as the limit $j \rightarrow 0$ with T finite. This means that one should be able to find a black-hole solution with *zero* source but finite temperature and with horizon exactly at ϕ_c . These flows are driven by the sub-leading term in the UV scalar field expansion, corresponding to a vev of the dual operator in the absence of a source. They will be discussed in detail in section 5.

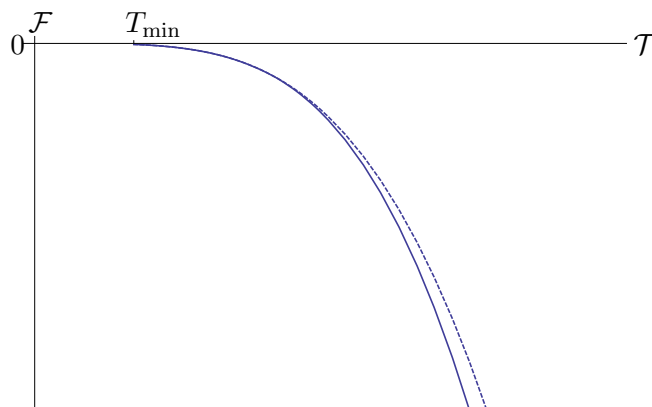


Figure 15. Free energy difference as a function of the (dimensionless) temperature between the two branches of black holes emanating from UV_2 with positive source (dashed and solid purple curves in figures 11 and 12).

Zero temperature limits. The limit $\mathcal{T} \rightarrow 0$ occurs when ϕ_h coincides with one of the IR fixed points (in which case we recover the two zero-temperature vacuum flows, skipping and non-skipping. Additionally, $\mathcal{T} \rightarrow 0$ as $\phi_h \rightarrow \phi_*$, which is somewhere in the middle between UV_2 and IR_2 . This may seem puzzling as there is no zero-temperature solution which ends at ϕ_* , since this does not correspond to an extremum of the potential. The puzzle can be resolved by tracking how the horizon approaches ϕ_* from both sides. This is represented in figure 16 in terms of the superpotentials of the various branches:

- approaching ϕ_* from the left ($\phi_h < \phi_*$), we have solutions starting at UV_2 with negative source, bouncing before IR_1 and ending close to ϕ_* . These are represented as the blue and violet flows in figure 16. The closer the horizon is to ϕ_* , the closer the bounce is to IR_1 .
- from the right ($\phi_h > \phi_*$), solutions start at UV_1 , skip IR_1 and end beyond ϕ_* . These are represented as the red and orange flows in figure 16. The closer the horizon is to ϕ_* , the closer the solution approaches IR_1 without touching it.

One can see that, *in the limit* $\phi_h \rightarrow \phi_*$ (represented in figure 17), both classes of solutions (starting from UV_1 and UV_2) actually reach IR_1 from both sides, and coincide with the ones represented in figure 6. These flows stop at IR_1 as the scale factor goes to zero there, and the solution is approaching asymptotically the Poincaré-AdS horizon. The remaining leftover piece (dashed black line in figure 17) starts from IR_1 and arrives to a horizon exactly at $\phi_h = \phi_*$.

However, for this last flow the point IR_1 is seen as a *UV fixed point*: recall that the superpotential always increases along the flow. These black holes therefore are thermal states in a different theory, the one for which IR_1 is a UV fixed point and have finite temperature *as defined in the UV theory sitting at IR_1* .

To summarise, in the limit $\phi \rightarrow \phi_*$, each black-hole solution from UV_1 and UV_2 splits into two disconnected solutions: a zero-temperature flow ending at IR_1 (red and purple

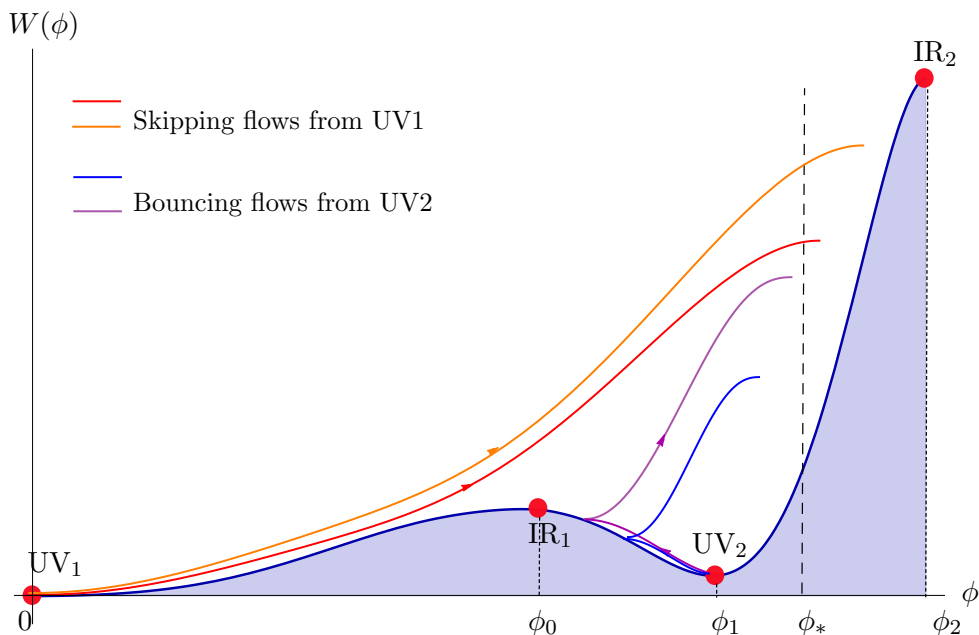


Figure 16. The approach to $\phi_h = \phi_*$ from both sides. The curves show the superpotentials of the corresponding flows. Here we display flows with different temperatures (longer flows correspond to lower temperatures).

curves in figure 17), and a finite temperature (vev-driven) flow starting in the UV from IR_1 and having its horizon at ϕ_* (dashed black curve in figure 17).

Like those ending at $\phi = \phi_c$, these black holes are also driven by a vev, and they will be discussed in more detail in section 5.

We make one last comment about an interesting feature of the solutions with horizon approaching ϕ_* : as can be observed in figure 9, although the solutions approaching the “gap” from the left and from the right are disconnected, it appears that the vev parameter C takes the same value on each side of the gap, i.e. $C(\phi_0) = C(\phi_*)$. This implies continuity of the free energy across the gap, since temperature and entropy approach zero at both ϕ_0 and ϕ_* . This can indeed be understood analytically. Indeed, consider one of the solutions starting from UV_1 in figure 16. As $\phi_h \rightarrow \phi_*^+$, we can regard it as composed by a flow W_1 which stops just above ϕ_0 plus a flow W_2 which starts just above ϕ_0 and reaches around ϕ_* . In the limit $\phi_h \rightarrow \phi_*$, the flow W_1 reduces to the non-skipping zero-temperature flow (red line in figure 17) while W_2 reduces to the flow from IR_1 to ϕ_* (dashed black line in figure 17). The free energy can be then written as a sum of the two contributions,

$$\mathcal{F}(\phi_h \rightarrow \phi_*^+) = \mathcal{F}_1 + \mathcal{F}_2 \tag{3.8}$$

By construction, in the limit we are considering, \mathcal{F}_1 approaches the free energy of the zero-temperature non-skipping solution with endpoint at ϕ_0 ,

$$\mathcal{F}_1 = \mathcal{F}(\phi_0). \tag{3.9}$$

As we will now show below, the contribution \mathcal{F}_2 is vanishingly small as $\phi_h \rightarrow \phi_*$. In fact, the flow W_2 can be also approached, in this limit, by the upper branch of the flow $W'_1 + W_2$

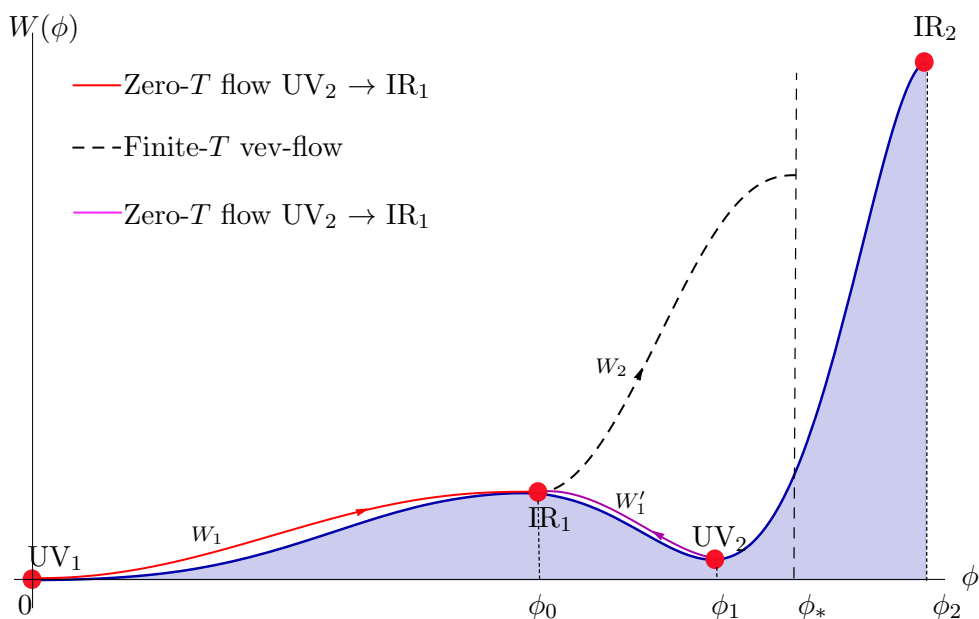


Figure 17. In the limit $\phi_h \rightarrow \phi_*$ the solutions reduce to the two zero-temperature flows ending at IR_1 , plus a finite-temperature flow starting at IR_1 in the UV and having its horizon at $\phi = \phi_*$ (dashed black curve). This solution has a finite temperature and entropy since it has a regular black-hole horizon.

in figure 17, which starts from UV_2 , and bounces very close to IR_1 : this is the limit of the blue and purple curves in figure 16 when the bounce point approaches IR_1 . We can view the branch W_2 after the bounce as a flow from a finite UV cut-off scale (with cut-off energy proportional to $e^{A(u_b)}$) to the horizon at $\phi_h \simeq \phi_*^-$. Its contribution to the total free energy of the solution is finite, and given by (see appendix B for details):

$$\frac{\mathcal{F}_2}{M_p^{d-1} V_{d-1}} = e^{dA(u_b)} W(\phi_b) \sqrt{f(u_b)} \tag{3.10}$$

where u_b is the radial coordinate of the bounce and ϕ_b is the corresponding value of the scalar field. Similarly, the bounce acts as an IR cut-off (with the same scale $e^{A(u_b)}$) for the lower branch W'_1 , connecting to UV_2 . Now, as $\phi_h \rightarrow \phi_*^-$, the bounce point ϕ_b approaches the IR fixed point ϕ_0 , the cut-off scale $e^{A(u_b)} \rightarrow 0$ and the contribution (3.10) to the free energy vanishes,

$$\mathcal{F}_2 \rightarrow 0, \quad \phi_h \rightarrow \phi_*. \tag{3.11}$$

Putting together equations (3.8), (3.9) and (3.11) we conclude that

$$\mathcal{F}(\phi_*) = \mathcal{F}(\phi_0). \tag{3.12}$$

4 Thermodynamics of bouncing RG flows

In this section we discuss another “exotic” kind of behaviour, which is unusual from the perturbative field theory standpoint, but which can be found in holographic RG flows [12, 17].

This consists in solutions where the flow inverts its direction at some $\phi = \phi_b$ (“bounce”). Although the corresponding β -functions are non-analytic at this point, $\beta(\phi) \sim \sqrt{|\phi - \phi_b|}$, the corresponding gravity solutions are regular and can be continued past the point u_b where the bounce takes place. Specifically, around $u = u_b$ the zero-temperature solutions have the expansion,

$$\begin{aligned} A(u) &= A_b^{(0)} + (u - u_b)A_b^{(1)} + O\left((u - u_b)^2\right), \\ \phi(u) &= \phi_b + \frac{V'(\phi_b)}{2}(u - u_b)^2 + O\left((u - u_b)^3\right), \end{aligned} \tag{4.1}$$

where $A_b^{(0)}, A_b^{(1)}$ are constants. From equation (4.1) it is clear that bounces can occur at any point in field space such that $V'(\phi_b) \neq 0$, in a such a way that ϕ has a maximum at ϕ_b if $V'(\phi_b) < 0$, and a minimum if $V'(\phi_b) > 0$.

4.1 Bouncing solutions at zero temperature

In [12], several examples were presented of vacuum holographic RG flow solutions exhibiting bounces. Some of them interpolate between a UV and an IR fixed point, and go through one or more bounces somewhere in the middle. In other examples, the potential had a single extremum (a UV maximum at $\phi = 0$ and the solution reaches $\phi \rightarrow \pm\infty$ after a bounce at a finite ϕ_b). The latter case corresponds to the zero-temperature solution in a model which was already considered at finite temperature⁷ in [17].

Bounces in the ground state solution occur when the dilaton potential is sufficiently steep. For example, if the potential behaves as $V(\phi) \propto \exp(\Gamma\phi)$ for large ϕ then one finds a bouncing zero T solution as the exponent is larger than a critical value $\Gamma > \Gamma_c$. To be concrete, we consider the potential

$$\begin{aligned} V(\phi) &= -\frac{d(d-1)}{\ell^2} \left(\cosh(\Gamma\phi) - \frac{1}{2}\phi^2\Gamma^2 \right) + \frac{1}{2\ell^2}\phi^2\Delta_-(\Delta_- - d) \\ &\quad - \frac{c}{\ell^2} \left(e^{-\frac{(\phi-\phi_0)^2}{2\sigma}} + e^{-\frac{(\phi+\phi_0)^2}{2\sigma}} + \phi^2 e^{-\frac{\phi_0^2}{2\sigma}} \left(\frac{1}{\sigma} - \frac{\phi_0^2}{\sigma^2} \right) - 2e^{-\frac{\phi_0^2}{2\sigma}} \right). \end{aligned} \tag{4.2}$$

This corresponds to a potential with a cosh and a quadratic term modified with two gaussian peaks with width σ and amplitude c added at $\phi = \phi_0$ and $\phi = -\phi_0$. Without the gaussian modification, ϕ would run from the UV conformal fixed point at $\phi = 0$, to the IR region at $\phi \rightarrow +\infty$ but the gaussians introduce an IR conformal fixed point at $\phi \approx \phi_0$ hence the IR behaviour is conformal. The quadratic terms in the potential above adjust the UV dimension of the perturbing operator to be $\Delta_+ = d - \Delta_-$. We plot this potential for a particular choice of parameters in figure 18, and we show the bouncing zero-temperature (vacuum state) solution in figure 19.

4.2 Bounces at finite temperature and the phase diagram

As in section 3, one has to determine T as a function of ϕ_h to describe all the black-hole solutions at a given T . Figure 20 shows a comparison of the bouncing vacuum solution, a non-bouncing finite T solution with $\phi_h = 0.7$ and a bouncing finite T solution with $\phi_h = 1.7$.

⁷The corresponding bouncing RG flow in the vacuum state has been also worked out but not discussed in [17].

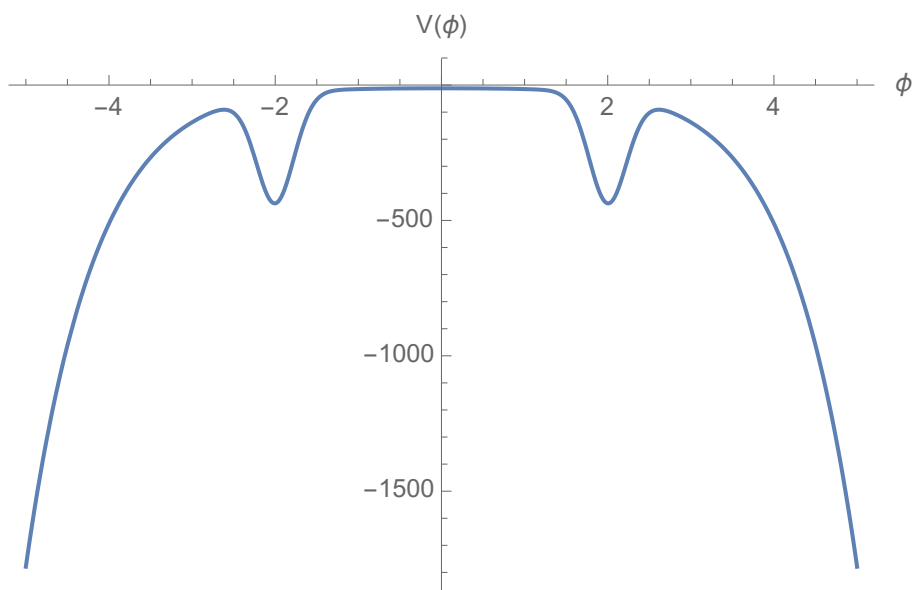


Figure 18. Example of a dilaton potential that leads to bouncing solutions. The potential is as in equation (4.2) with a choice of parameters $d = 4$, $\Delta_- = 1.7$, $\phi_0 = 4$, $\Gamma = 2/\sqrt{3}$, $c = 400$, $\sigma = 1/20$. The UV fixed point corresponds to $\phi = 0$ and the IR fixed point corresponds to $\phi \approx 2.00566$.

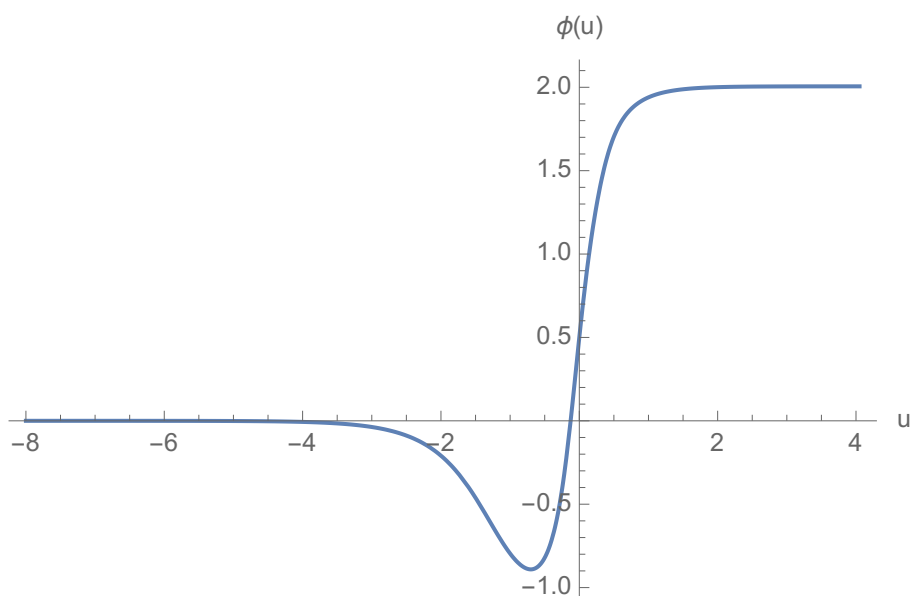


Figure 19. Bouncing zero-temperature solution for the potential shown in figure 18.

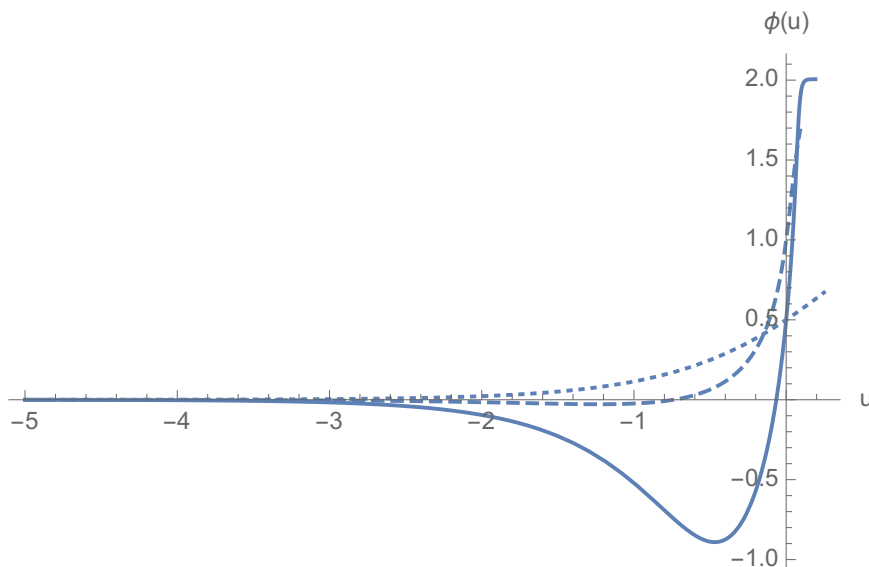


Figure 20. Comparison of the vacuum solution (solid), a bouncing finite T solution with $\phi_h = 1.7$ (dashed) and a non-bouncing finite T solution with $\phi_h = 0.7$ (dotted) that arise from the potential shown in figure 18.

Because there can be more than one ϕ_h corresponding to the same value of T , one needs the free energy to distinguish between these multiple black-holes and to determine the dominant one with the lowest free energy.

Calculating T as a function of ϕ_h reveals the different black-hole solutions at a given temperature. This can be obtained either by constructing the geometry and using equation (2.5) or more directly from the potential using the method of scalar variables explained in appendix C, see equation (C.12). Following the latter method we obtain the function $T(\phi_h)$ shown in figure 21.

There are four separate intervals of ϕ_h that we label from I to IV in this figure. First, we note the symmetry $T(\phi_h) = T(-\phi_h)$. It is not hard to see that this is an immediate consequence of the symmetry $V(\phi) = V(-\phi)$ of the dilaton potential. Moreover, there exists a critical value of ϕ_h , that reads $\phi_c = 1.507$ for the choice of parameters in figure 18, above which the black-hole solution becomes bouncing. This is demonstrated in figure 20 where we show a bouncing type black-hole for $\phi_h = 1.7$ and a non-bouncing type for $\phi_h = 0.7$. These two solutions belong to regions I and IV in figure 21. It is important to realise that these two regions I and IV *do not belong to the same theory*: as $\phi = j \ell^{\Delta-} e^{\Delta- u/\ell} + \dots$ near the boundary, the solutions in region IV have $j > 0$ whereas the solutions in region I have $j < 0$ because of the bounce. Therefore the solutions in regions I and IV belong to different boundary theories.

Similarly, the solutions in region II have $j < 0$ and region III have $j > 0$. In the following, we will choose $|j| = 1$ with no loss of generality (with this choice, the dimensionless temperature parameter \mathcal{T} defined in equation (2.14) coincides with the temperature T). To simplify the presentation, we further define two reference scales T_{ref} and F_{ref} are defined in (C.16).

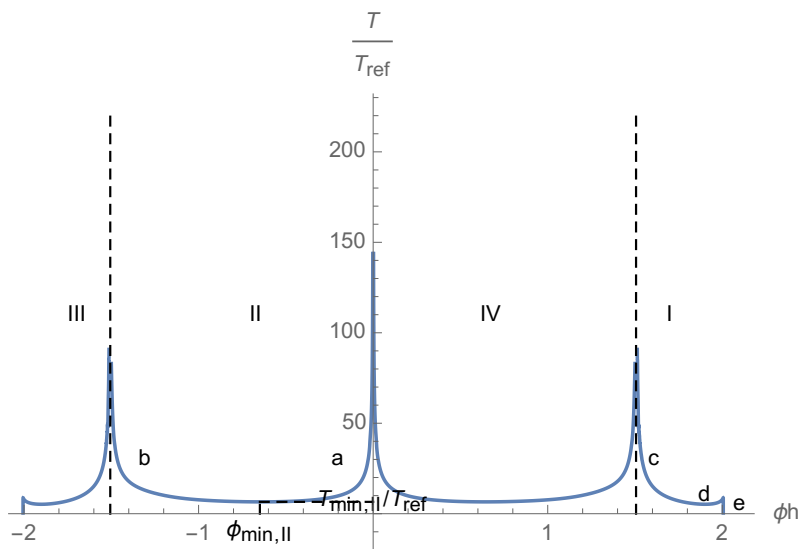


Figure 21. T as a function of ϕ_h for a particular choice of the reference scale T_{ref} .

The symmetry of $V(\phi)$ under parity implies that we may only consider the solutions in regions I (bouncing) and II (non-bouncing). The physics of regions III and IV are the same except for the sign of j .

4.2.1 Region I: bouncing black-holes

First consider region I. As shown in figure 21, this region is further divided into three subregions, between $\phi_c \leq \phi_h < \phi_{\text{min,I}}$, $\phi_{\text{min,I}} \leq \phi_h < \phi_{\text{max}}$ and $\phi_{\text{max}} \leq \phi_h < \phi_{IR}$, labelled on the figure as “c”, “d” and “e” respectively, where for our choices of parameters we have $\phi_{\text{min,I}} = 1.901$ and $\phi_{\text{max}} = 2.005$. The corresponding temperatures read $T_{\text{min,I}} = 5.209 T_{\text{ref}}$ and $T_{\text{max}} = 8.685 T_{\text{ref}}$.

As clear from figure 22 there are three different black-hole solutions in this region at a given temperature between $T_{\text{min,I}}$ and T_{max} .

Below $T_{\text{min,I}}$ and above T_{max} there is a single type of black-hole solution. The free energy of these solutions can be calculated following the method described in appendix C. The result is shown in figure 23. There is a first order phase transition at $T_c \approx 6T_{\text{ref}}$. This conclusion however will change once we consider the solutions in region II.

Another point to note is the free energy of the branch between $T = 0$ and $T = T_{\text{max}}$, as shown figure 21 (b). The free energy of this branch is almost constant (with almost vanishing entropy) as shown in figure 23. In the limit $T \rightarrow 0$ this branch should smoothly turn into the vacuum solution. Indeed, the limiting value of the free energy as $T \rightarrow 0$ coincides precisely with the free energy of the vacuum solution.

4.2.2 Region II: non-bouncing black-holes

In this region there are three black-hole solutions. One observes at least two branches of solutions in figure 21, for $0 \geq \phi_h > \phi_{\text{min,II}}$ and for $\phi_{\text{min,II}} \geq \phi_h > -\phi_c$ where $\phi_{\text{min,II}} = -0.649$. The corresponding temperature reads $T_{\text{min,II}} = 6.545 T_{\text{ref}}$.

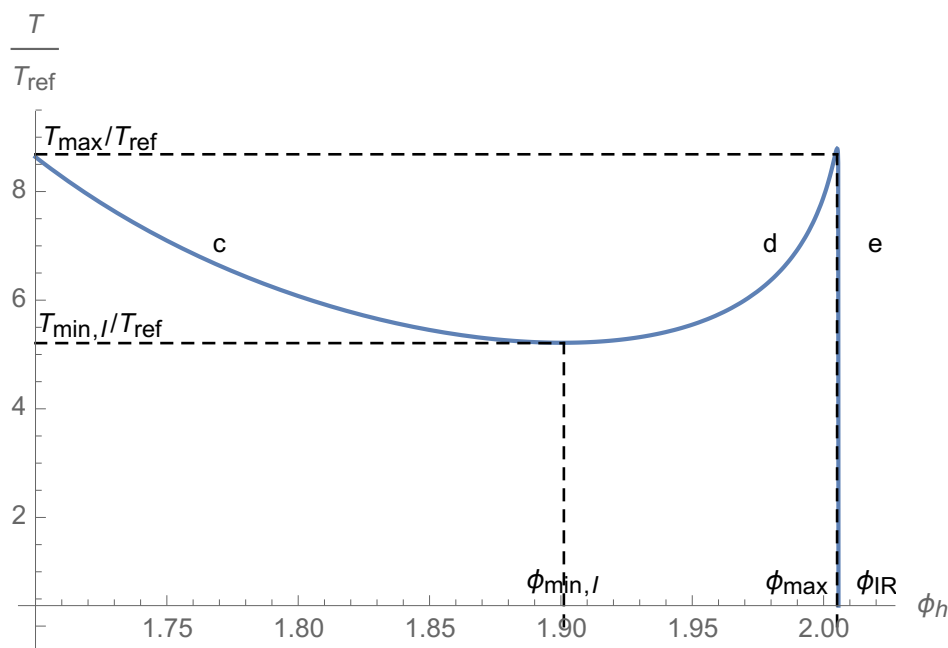


Figure 22. Blowing up the figure 21 at very small temperatures reveals existence of there different blackhole solutions for T fixed between a $T_{\min,I}$ and a T_{\max} .

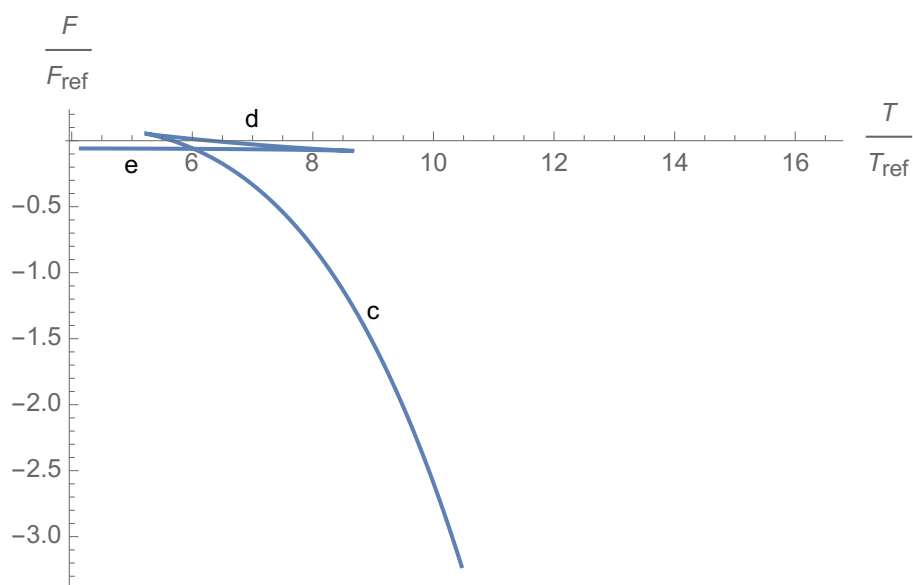


Figure 23. Free energy of bouncing type solutions in region I as a function of T that follows from the potential in figure 18. Latin letters label the different branches shown in figure 21.

However, there is a hidden third branch in this region, that is unveiled when one plots the entropy as a function of T . This is shown in figure 24 (a) and (b). Because of the shape of the function $s(T)$ near the minimum, the aforementioned second branch for $\phi_{\min, \text{II}} \geq \phi_h > -\phi_c$ is further divided into two, as $T_{\min, \text{II}} < T < T_s$ and $T_s < T < \infty$ where T_s corresponds to the minimum of s in figure 24 (b), with the value $T_s = 6.560 T_{\text{ref}}$. The corresponding value of ϕ_h is $\phi_s = -0.691$. Therefore, we need to divide the branch “b” in figure 21 further into two: the branch for $\phi_{\min, \text{II}} \geq \phi_h > \phi_s$ and for $\phi_s \geq \phi_h > -\phi_c$.

The need for this further division of branches becomes more apparent upon consideration of the speed of sound in the plasma. The speed of sound can be expressed in terms of entropy and temperature as

$$c_s^2 = \frac{s}{T \frac{ds}{dT}}. \tag{4.3}$$

As s and T are positive definite everywhere, and that dS/dT vanishes (diverges) at $T = T_s$ ($T = T_{\min, \text{II}}$) one finds that c_s^2 ranges between 0^- and $-\infty$ between $T_{\min, \text{II}} < T < T_s$. The reason that $c_s^2 \rightarrow -\infty$ as $\phi_h \rightarrow \phi_s^+$ is because $ds/dT \rightarrow 0^-$ there. On the other hand $c_s^2 \rightarrow +\infty$ as $\phi_h \rightarrow \phi_s^-$, thus c_s^2 of this solution jumps by an infinite amount at $T = T_s$. Therefore one has to characterise the solutions for $T < T_s$ and $T > T_s$ differently. That $ds/dT < 0$ between $T_{\min, \text{II}} < T < T_s$ implies that this branch is thermodynamically unstable: it has negative definite specific heat⁸ per unit volume $C_v = T ds/dT$.

To complete the discussion of the speed of sound, we find that c_s^2 ranges between⁹ $1/3$ and 0^+ in branch “a” on figure 21 with $0 \geq \phi_h > \phi_{\min, \text{II}}$ and it ranges between $+\infty$ and some positive value (which should be determined by precise numerics near $-\phi_c$) in the branch with $\phi_s \geq \phi_h > -\phi_c$. The fact that the speed of sound can exceed 1 in this branch most probably means that this branch is dynamically unstable, in analogy with the example studied in [17].

Comparison of the free energies of the blackhole branches that belong to region I and II are shown in figure 26. One observes that the solutions in region II are always dominant. We observe that, as T is increased first there is a first order transition between the blackhole branches “e” and “c” in region II at $T_c/T_{\text{ref}} = 6.1$ and then there seems to be a jump in the free energy at $T = T_{\min, \text{II}} = 6.545 T_{\text{ref}}$ from branch “c” to branch “b”. This jump in the free energy may indicate that the finite temperature extension of bouncing holographic RG flows is ill defined.

Finally in figure 27 we plot and compare the vacuum expectation values of the scalar operator $\langle \mathcal{O} \rangle$ on the blackhole branches in regions I and II.

⁸Therefore there exists a very small thermodynamically unstable region on the (generically thermodynamically stable) upper branch of solutions shown in figure 25 between the cusp at $T = T_{\min, \text{II}}$ and T_s . At T_s C_v changes sign and becomes positive for $T > T_s$. This does not imply any non-analyticity for the free energy at T_s however.

⁹ $1/3$ because the plasma becomes conformal as $\phi_h \rightarrow 0$ and 0^+ as $\phi_h \rightarrow \phi_s$ because dS/dT diverges there.

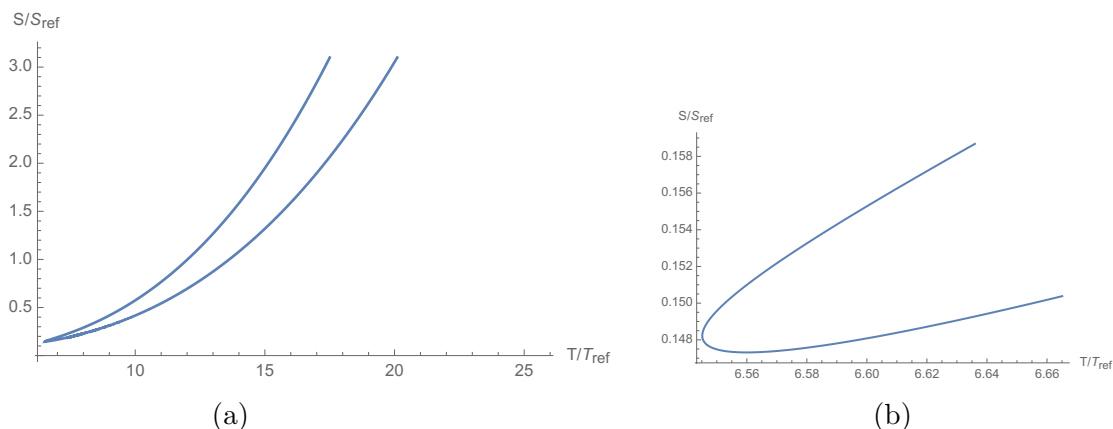


Figure 24. (a) S as a function of T in region II. (b) Blow up of the figure around the minimum of S .

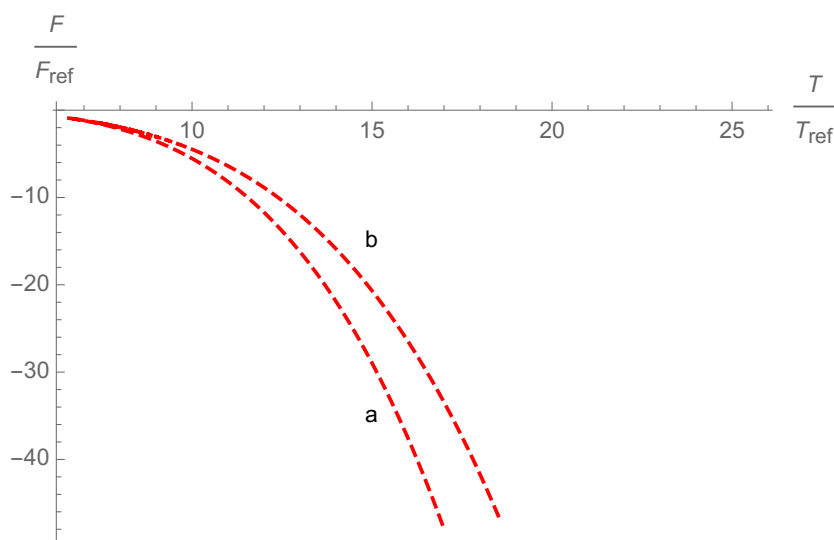


Figure 25. Free energy of non-bouncing type solutions in region II as a function of T that follows from the potential in figure 18. Latin letters label the different branches shown in figure 21.

5 Sourceless black holes

As we have seen in the previous two sections, there are some special values of the horizon position in field space, ϕ_h , which correspond to singular limits:

- Approaching the values ϕ_c which separates between bouncing and non-bouncing solutions, the dimensionless thermodynamic quantities diverge;
- In the case considered in section 2, there is a special point ϕ_* where (at fixed source) the temperature goes to zero, though this value does not obviously correspond to the endpoint of any vacuum RG flow.

In this section we clarify the meaning of these special points. As we will see, they are new families of solutions, which correspond to flows driven by the vev of a relevant (in the

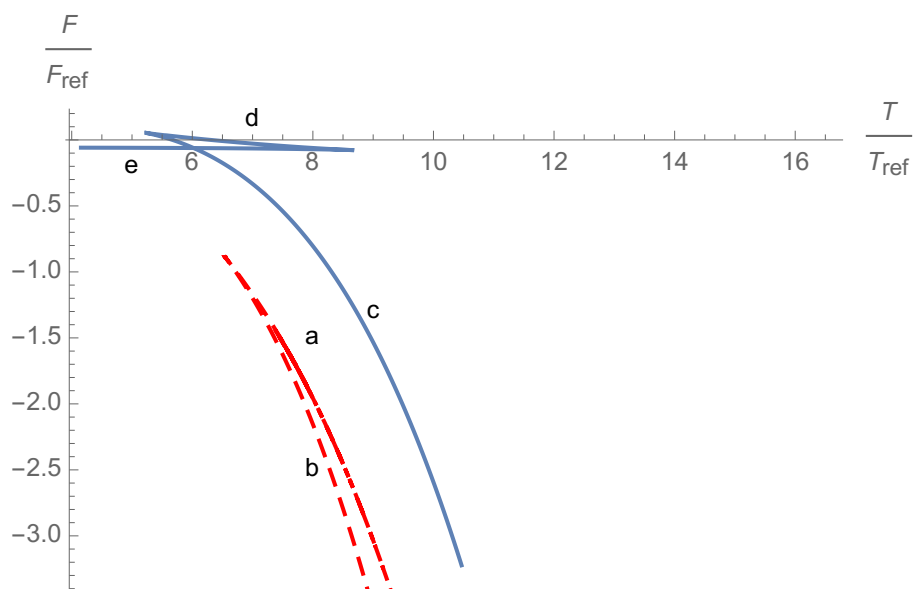


Figure 26. Comparison of the free energies of the solutions that belong to region I (blue, solid) and II (red, dashed).

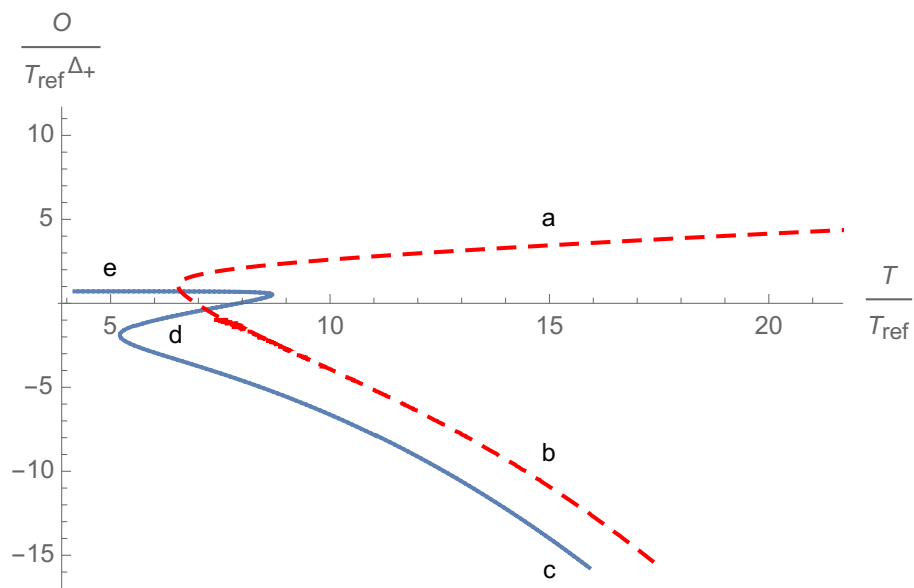


Figure 27. Comparison of the VEVs of the scalar operator on the various black hole branches in region I (blue, solid) and region II (red, dashed). Latin letters label the separate branches shown in figure 21.

case of ϕ_c) or an irrelevant (in the case of ϕ_*) operator. Finally, we will analyse the fate, at finite temperature, of solutions in special (fine-tuned) theories which at $T = 0$ admit regular sourceless flows from a minimum to a minimum of the bulk potential, an example of which was presented in [12].

5.1 Thermodynamics of vev-driven flows

It is useful to analyse vev-driven flows in terms of the superpotential formalism developed in appendix A. In this language, a sourceless flow is a solution of the first order equations governed by the superpotential of the type W^+ , of the form given in equation (2.20),

$$\dot{A} = -\frac{1}{2(d-1)}W^+(\phi), \quad \dot{\phi} = \frac{dW^+}{d\phi}. \tag{5.1}$$

Unlike W^- , which contains the vev-related constant $C(\phi_h)$ as an integration constant (see equation (2.18)), the W^+ solution is unique and does not admit continuous deformations. Therefore the solution, when expressed in terms of the scalar field as an independent variable in the superpotential formalism, is completely fixed, including the horizon position. The integration constant of the first order equation, ϕ_+ , is the only remaining integration constant of the solution, (recall that the integration constant in A is fixed by the requirement that the boundary metric is $\eta_{\mu\nu}$ without any scaling factor). Therefore, vev-driven flows form a one-parameter family of solutions, parametrised by either T or ϕ_+ . This is to be contrasted with source-driven flows, which are parametrised by the two independent data (T, ϕ_-) .

To make this more explicit, we can now repeat the scaling argument of section A.4, which shows that solutions with different values of ϕ_+ are generated by the transformation

$$\phi_+ \rightarrow e^{\Delta_+ v} \phi_+, \quad T \rightarrow T e^v, \tag{5.2}$$

which is the analog of the symmetry (A.29)–(A.31) in the case of non-zero source. This transformation leaves ϕ_+/T^{Δ_+} invariant, therefore this quantity must have the same value for all vev-driven flows.¹⁰ We conclude that, for vev-driven black holes attached to the same UV fixed point, the vev and the temperature are not independent parameters but they must obey

$$\langle O \rangle = \bar{\mathcal{O}} T^{\Delta_+} \tag{5.3}$$

where $\bar{\mathcal{O}}$ is a fixed, temperature-independent, dimensionless constant.

Finally, it is useful to repeat the calculation of the free energy in the case of sourceless flows. The calculation follows the same steps detailed in appendix B, except that instead of the solution W^- (equation (2.18)) we have to use the superpotential of the type W^+ given in equation (2.20). The crucial difference in this case is the absence of the $C\phi^{d/\Delta_-}$, which contributed the second term in the free energy (2.23): in this case there is no finite term coming from the sub-leading non-analytic part of the superpotential.

¹⁰This applies for vev-driven flows which are connected to the same UV fixed point. If there are multiple UVs, there can be several vev-driven solutions, but there is at most one for each UV fixed point.

Therefore, for sourceless flows governed by the superpotential W^+ , we simply have

$$\mathcal{F}_{\text{vev}} = -\frac{T s}{d} V_{d-1}. \quad (5.4)$$

This implies that vev-driven flows have conformal thermodynamics: integrating the relation $s = -(1/V_{d-1})\partial\mathcal{F}/\partial T$ gives

$$s = \bar{\sigma} T^{(d-1)}, \quad \mathcal{F}_{\text{vev}} = -\frac{\bar{\sigma}}{d} V_{d-1} T^d, \quad (5.5)$$

where $\bar{\sigma}$ is a fixed constant which only depends on the details of the bulk potential.

The fact that sourceless black holes display conformal thermodynamics is expected since, for zero source, conformal invariance is always softly broken. An alternative derivation of this result, based on the effective potential, will be presented in the next subsection.

5.2 Relevant vev flows

First, we consider the black holes for which the dimensionless temperature (defined in (2.14), $\mathcal{T} \rightarrow \infty$ as the horizon approaches a critical value ϕ_c , which is in between extrema of the potential. As explained in the previous sections, the value ϕ_c separates between bouncing and non-bouncing solutions, which have opposite values of the source. This change of sign of the source is perceived as a divergence in the dimensionless temperature \mathcal{T} , which can be interpreted in two different ways depending how we approach ϕ_c in the (T, ϕ_-) space:

1. If we consider the theory with fixed, finite source, then none of the solutions can have a horizon at $\phi_h = \phi_c$. As we approach this limit, the temperature $T \rightarrow \infty$ and so do the entropy and the free energy. There is no regular solution in this limit.
2. On the other hand, we can approach the $\phi_h = \phi_c$ by keeping T fixed and sending $\phi_- \rightarrow 0$. Then we obtain *regular* black-hole solutions with finite free energy, a horizon exactly at ϕ_c , and zero source. For these black holes, the flow is driven instead by the vev of the (relevant) operator dual to ϕ .

In the rest of this subsection, we will examine further these critical black-hole solutions with horizon at ϕ_c .

We start by noting that the limit in which the flow becomes vev-driven corresponds to the following scaling limit in the parameters entering the solution (see equation (2.19))

$$\phi_- \rightarrow 0, \quad C \rightarrow +\infty, \quad \frac{d}{\Delta_-} \frac{C|\phi_-|^{\Delta_+/\Delta_-} \text{sign}(\phi_-)}{(d-2\Delta_-)} = \phi_+^c \text{ fixed} \quad (5.6)$$

In this limit, the vev $\langle \mathcal{O} \rangle = (d-2\Delta_-)\phi_+^c$ remains finite. As we will see shortly, the value ϕ_+^c is not free but it is determined by the temperature.

The most transparent way to understand the thermodynamics of the vev-driven solutions ending at ϕ_c is to use the effective potential derived in section 2.5. From equations (2.31) and (2.33), we see that solutions with $\phi_- = 0$ correspond to extrema of the effective potential (2.33), for which

$$\mathcal{V}'(\mathcal{O}) = 0 \quad \Rightarrow \quad \mathcal{O} = \mathcal{O}^c \quad (5.7)$$

where $\mathcal{O} = \langle \mathcal{O} \rangle / T^{\Delta_+}$ and \mathcal{O}^c is a constant. This equation therefore fixes the vev in terms of the temperature,

$$\langle \mathcal{O} \rangle = \mathcal{O}^c T^{\Delta_+} \tag{5.8}$$

Therefore, there is a one-parameter family of black holes, parametrised by T , all having a horizon at $\phi = \phi_c$, zero source, and vev given by equation (5.8).

On-shell, since the source is zero, the effective potential is the same as the free energy, and both are given by

$$\mathcal{F} = V_{\text{eff}}(T, \mathcal{O}^c) = \mathcal{V}_c T^d \tag{5.9}$$

where $\mathcal{V}_c = \mathcal{V}(\mathcal{O}^c)$ is a temperature-independent constant. We have recovered the general fact, discussed in subsection 5.1, that this family of black holes displays conformal thermodynamics, equation (5.5). Comparing equation (5.9) with (5.5) we can read-off

$$\mathcal{V}_c = -\frac{\sigma^c}{d} V_{d-1} \tag{5.10}$$

where σ^c is the (fixed) ratio s/T^{d-1} .

Notice that, for fixed T and $j = 0$, the theory has another black hole solution with the same UV asymptotics: it is the AdS-Schwarzschild black hole with constant scalar field $\phi = \phi_{\text{UV}}$, sitting at the maximum of the potential. This solution also has conformal thermodynamics, with

$$\mathcal{F}_{\text{AdSS}} = -\frac{\sigma_{\text{conf}}}{d} V_{d-1} T^d, \quad \sigma_{\text{conf}} = \frac{(4\pi)^d}{d^{d-1}} (M_p \ell_{\text{UV}})^{d-1} \tag{5.11}$$

where we have expressed σ_{conf} using equation (2.7). The question then arises, which of these two solutions is thermodynamically favored. Because of the simple scaling behaviour of the free energy, the answer is the same at all $T \neq 0$, and it boils down to comparing the values of σ_{conf} and σ^c . A numerical computation shows that, in the particular models we considered here, $\sigma^c < \sigma_{\text{conf}}$, meaning that the AdS-Schwarzschild solution is the dominant one. For relevant vev flows ending at ϕ_c in the model in section 3, we find $\sigma^c = 0.95\sigma_{\text{conf}}$; we come to the same conclusion for the corresponding solutions in the model considered in section 4, where we find $\sigma_c = 0.635\sigma_{\text{conf}}$.

Finally, we can go slightly off-shell and analyse the behaviour of the effective potential around the critical value \mathcal{O}^c . Using the scaling property (5.6) it is easy to show that, as $\mathcal{T} \rightarrow \infty$, the quantities $\gamma(\mathcal{T})$ and $\sigma(\mathcal{T})$ in equation (2.29) behave as

$$\sigma(\mathcal{T}) \rightarrow \sigma^c, \quad \gamma(\mathcal{T}) \simeq \mathcal{T}^{-\Delta_-} \rightarrow 0. \tag{5.12}$$

Performing the Legendre transform of equation (2.28) explicitly close to $\phi_- = 0$ we find, as $\mathcal{O} \rightarrow \mathcal{O}^c$,

$$\mathcal{T} \simeq -\frac{k}{\mathcal{O} - \mathcal{O}^c}, \quad \mathcal{V}(\mathcal{O}) \simeq \mathcal{V}_c + \frac{|\mathcal{O} - \mathcal{O}^c|^{\Delta_- + 1}}{(\Delta_- + 1)k^{\Delta_-}}, \tag{5.13}$$

where k is a constant.

5.3 Fake zero- T vacua and irrelevant flows

We now turn to another kind of special flows, which correspond to black holes ending at the point ϕ_* found in section 3 and represented in figure 17. As we have explained at the end of section 3, these black holes correspond to flows starting from the point IR_1 , located at $\phi = \phi_0$ in figure 17, which in this case plays the role of a UV fixed point.

Since the UV is at a minimum of the potential, the conformal dimension of the deforming operator is $\Delta_+ > d$, the operator is irrelevant, and a source term is not allowed. Therefore, the solutions with horizon at ϕ_* are driven by a vev, as those discussed in the previous section. This is consistent with the fact that, if we start from the UV at ϕ_0 , there is a *unique* value ϕ_* the scalar can take at the horizon: it is fixed by the unique solution W^+ starting at $\phi = \phi_0$.

By the general discussion at the beginning of this section, there is a one-parameter family of black holes, labeled by the temperature T , for which the vev given by

$$\langle O \rangle = \mathcal{O}^* T^{\Delta_+}, \tag{5.14}$$

where \mathcal{O}^* is a constant.

The free energy is given by the conformal result

$$\mathcal{F} = -\frac{\sigma^*}{d} V_{d-1} T^d \tag{5.15}$$

where $\sigma^* = s/T^{(d-1)}$ is temperature-independent

Notice that the free energy is defined by renormalising with respect to IR_2 , i.e. the counter-term must be chosen to be

$$S_{\text{ct}} = M_p^{d-1} \int d^d x \sqrt{g} \frac{2(d-1)}{\ell^2(\phi_0)}, \quad \ell(\phi_0) \equiv \sqrt{-\frac{d(d-1)}{V(\phi_0)}}. \tag{5.16}$$

Since they have a different UV boundary condition (and different counter-terms) than those connecting to either UV_1 or UV_2 , the black holes ending at ϕ_* belong to a boundary theory different from the ones considered in section 3 and in the previous subsection. Therefore, unlike the case considered in the previous section, we cannot describe the free energy in terms of an effective potential at criticality, since there is no well-defined conjugate variable to \mathcal{O} which we can use to define a Legendre transform.

It is instructive to understand why these solutions seem to arise as a zero-temperature limit. As we have seen in section 3, when measuring the temperature as defined in UV_1 or UV_2 , the free energy receives a vanishing contribution from the part of the solution which connects ϕ_0 to ϕ_* , and the limit $\phi \rightarrow \phi_*$ looks like a zero-temperature limit. This is due to the fact that, from the point of view of, say, UV_1 , any solution starting from IR_1 is seen as describing the far infrared. Therefore, any finite temperature as measured in units of IR_1 will be rescaled to zero in units of UV_1 .

Notice that there is no sense in “glueing together” the two solutions composed of the flow from 0 to ϕ_0 and the flow to ϕ_0 to ϕ_* to obtain a new, exotic, zero-temperature solution: indeed, the vev-driven solution from ϕ_0 to ϕ_* reaches an asymptotically AdS UV

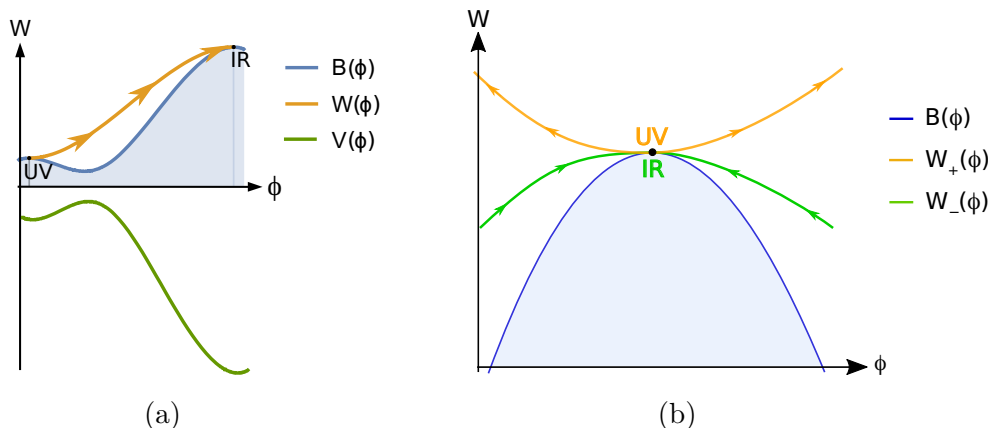


Figure 28. A minimum to minimum flow is displayed. The left figure (a) shows a schematic picture of the superpotential (yellow curve) and the bulk potential (green curve). The curve $B(\phi)$ is defined via equation (A.18) and is the boundary of the forbidden region below which the zero-temperature superpotential cannot go. The right figure (b) shows a sketch of solutions of the superpotential equation with critical point at a local minimum of the potential (local maximum of $B(\phi)$). The $W_+(\phi)$ and $W_-(\phi)$ solutions correspond each to two asymptotically AdS geometries in the UV and IR, respectively. The corresponding geometries are not connected, as each flow stops (or starts) at the fixed point. The arrows on $W_-(\phi)$ and $W_+(\phi)$ indicate the direction of the holographic RG flow.

boundary as $\phi \rightarrow \phi_0$, where $e^A \rightarrow +\infty$. This geometry is locally geodesically complete, and it cannot be glued across the horizon of the flow reaching ϕ_0 in the IR (where $e^A \rightarrow 0$).

Finally, a numerical computation of σ^* shows that, also in this case, the free energy (5.15) is larger than the free energy of the AdS-Schwarzschild solution of the same temperature and constant scalar field $\phi = \phi_0$. The latter is therefore the dominant solution at any temperature for $j = 0$.

5.4 Minimum-to-minimum irrelevant flows

In this section, we consider the finite-temperature generalisation of the flows connecting two *minima* of the potential (one serving as a UV fixed point, the other as an IR fixed point). These flows, discussed in [12], are driven by the vev of an *irrelevant* operator, in contrast to the zero-temperature solutions discussed in the sections 3 and 4, for which the operator was always relevant. They are shown schematically in terms of the associated superpotential in figure 28.

As we discuss below, at zero temperature, these theories display a moduli space of vacua.¹¹ We will see that this is lifted when temperature is turned on, and that the only point in the moduli space that is left is the AdS black hole with constant scalar field profile.

Minimum-to-minimum flows are interesting because they provide the only regular zero-temperature flows in this setup which are driven by the vacuum expectation value of an irrelevant operator. Such flows arise from spontaneous breaking of scale invariance of

¹¹A field theory example of such a flow is the flow driven by baryon vevs in the baryonic branch of N=1 sQCD.

the boundary theory: the source J of the UV theory operator O is zero as it is non-renormalisable and, yet, this operator acquires a non-zero VEV which is related to the asymptotic behaviour of $\phi(u)$ through

$$\phi(u) = \phi_+ \ell^{\Delta_+} e^{\Delta_+ u} + \dots, \quad \langle O \rangle = (M_p \ell)^{d-1} (d - 2\Delta_-) \phi_+. \quad (5.17)$$

Equation (5.17) is valid for both zero and finite temperature. For generic potentials and at zero temperature, the bulk geometries in the class (5.17) are singular. However, for special potentials, it is possible to make the flow reach a second minimum of V in the IR, providing a regular solution with IR AdS asymptotics [12].

The first-order or superpotential formalism of appendix A provides a single description of *all* zero-temperature flows of the form (5.17) as well as an easy implementation of the regularity condition. The flows of the form (5.17) with non-zero $\langle O \rangle$ which start from a given minimum of V (seen as a UV fixed point) are associated with a *unique* superpotential of the type $W^+(\phi)$, with an asymptotic expansion of the form (2.20). This is the yellow curve in figure 28 (b)), where, for comparison, we also displayed the W_- -type solution arriving at the same minimum for which this point is seen as an IR fixed point. The blue curve in figure 28, $B(\phi)$, is defined through equation (A.18) and it is the lower bound on W at zero temperature: the shaded blue region is not allowed, as a consequence of the superpotential equation (A.2) with $f \equiv 1$ (see appendix A.2 for more details).

At $T = 0$, these models display a moduli space of vacua, parametrised by ϕ_+ . The reason is that the superpotential which describes these flows is of the type W^+ which, as we explained in subsection 5.1, does not contribute a finite part $\sim C\phi^{d/\Delta_-}$ to the renormalised on-shell action. Equation (2.27) then implies that the zero-temperature free energy $\mathcal{F} = 0$ for any value $\langle O \rangle$.

This one-parameter degeneracy of vacua is continuously connected¹² to the AdS vacuum of the unbroken theory, which corresponds to $\langle O \rangle = 0$ and to a constant scalar field profile and also has vanishing free energy.¹³

At finite temperature, as we have seen in subsection 5.1, we generically expect for vev-driven black holes a relation between $\langle O \rangle$ and T of the form (5.8). This means that at any fixed temperature at most one solution is expected, and the moduli space is lifted. Moreover, taking the $T \rightarrow 0$ limit we only obtain an AdS solution with $\langle O \rangle = 0$ i.e. these black holes, if they exist, are connected only to the constant- ϕ solution and to no other solution in the moduli space.

The previous considerations suggest that the solutions with non-zero $\langle O \rangle$ and non-trivial scalar field profile do not have any finite-temperature generalisation: to obtain a finite $\langle O \rangle$ at $T = 0$ one would need the constant $\bar{\mathcal{O}}$ in equation (5.3) to be infinite, in which

¹²The limit $\phi_+ \rightarrow 0$ is not uniform however: for any non-zero ϕ_+ , eventually the scalar field reaches its fixed IR value, $\phi(u = +\infty)$, which is different from the UV value $\phi(u = -\infty)$. As $\phi_+ \rightarrow 0$, a significant departure from the UV value ϕ_{UV} happens for larger and larger u : the domain wall “moves to infinity” leaving the UV-AdS solution at any finite u .

¹³More precisely, the free energy is zero in our renormalisation scheme, in which we have chosen $C_{ct} = 0$ in the counter-term action, see equation (B.11). In a more general scheme \mathcal{F} will be a non-zero, $\langle O \rangle$ -independent constant.

case the finite temperature VEV would diverge. Since the W^+ -type potential admits no continuous deformations close to the UV [12, 34], there cannot be regular black holes with $\langle O \rangle = 0$ starting from the same UV minimum as the zero-temperature regular flow. The reason is that there exists a W^+ regular solution for $f = 1$ and therefore any other regular solution with non-trivial f will “miss” the UV fixed point and flow somewhere else.

We will now move to a concrete example in which we find that these expectations are correct: in the theory that admits a regular minimum-to-minimum flow at $T = 0$, there are no black-hole solutions which start from the same UV.

To be concrete, we consider the model presented in reference [12], which we will subsequently study at finite temperature. In d boundary dimensions, the following family of superpotentials parametrised by k , v and W_* ,

$$W(\phi) = kv^2 \left(\frac{\phi}{v}\right) \left[1 - \frac{1}{3} \left(\frac{\phi}{v}\right)^2\right] + W_* \tag{5.18}$$

allows for very simple kink scalar field profiles:

$$\phi(u) = v \tanh(k(u - u_*)). \tag{5.19}$$

Each superpotential in (5.18) solves (A.2) with $f(\phi) = 1$ and with the potential

$$V(\phi) = \frac{(kv)^2}{2} \left[1 - \left(\frac{\phi}{v}\right)^2\right]^2 - \frac{d}{4(d-1)} \left\{ kv^2 \left(\frac{\phi}{v}\right) \left[1 - \frac{1}{3} \left(\frac{\phi}{v}\right)^2\right] + W_* \right\}^2. \tag{5.20}$$

The flows of the form (5.19) interpolate between the extrema of $V(\phi)$ at $\phi = -v$ and $\phi = v$. A region of parameter space which includes the point

$$v = k = 1, \quad d = 4 \quad \text{and} \quad W_* = 1.8 \tag{5.21}$$

is such that the extrema of $V(\phi)$ at $\phi = \pm v$ are both local minima. The potential (5.20) is not everywhere negative, however it is negative between $-v$ and v for the values (5.21) of the parameters. As we argued in [12], the sign of $V(\phi)$ outside the range of definition of a superpotential $W(\phi)$ does not affect the existence of this kind of flow because the equations of motion are local in field space. In this particular model, the maximum of V in between the two minima violates the BF bound, so it does not provide a consistent UV for the theory. This is inconsequential as far as the minimum-to-minimum flow is concerned.

At finite temperature, we need to solve the system (A.2)–(A.3). Contrary to the zero-temperature case where the bound (A.18) holds as long as the potential is strictly negative, at finite temperature even strictly negative potentials allow a positive $W(\phi)$ to cross into the forbidden region and to be arbitrarily small, as follows for example from equation (A.19).

To scan for black-hole solution numerically, we choose a set of values of horizon position ϕ_h in the interval $(-v, v)$ and solve numerically the system (A.2)–(A.3) providing boundary conditions at $\phi_h + \delta\phi$ using the relations (A.8). What we find is that no solution reaches the would be UV point $\phi = -v$, but all of them bounce at some larger ϕ . The

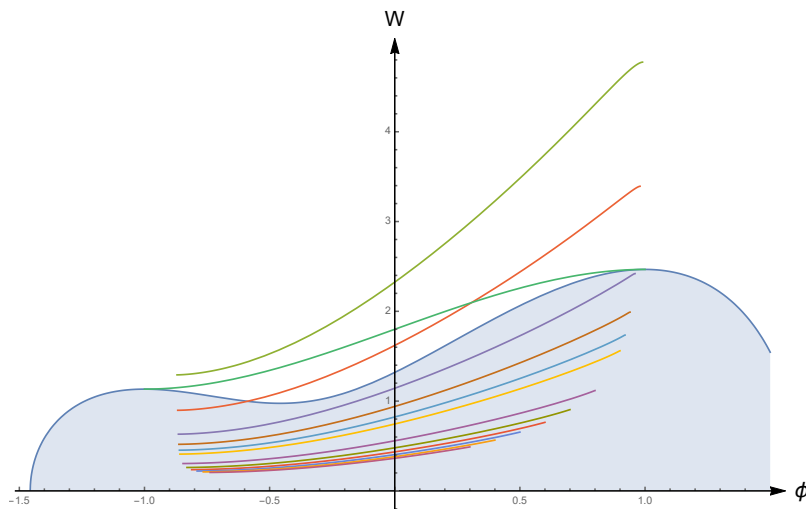


Figure 29. The blue curve bounding the shaded region represents the curve $B(\phi)$ which bounds superpotentials at zero temperature (see (A.18)). It is presented here to locate the extrema of the potential between which the zero temperature solution flows. The green curve which starts at -1 with value $B(-1)$ and extends until $\phi = 1$ where its value is $B(1)$ corresponds to the zero-temperature solution. The remaining curves correspond to finite-temperature superpotentials solving (A.2) and (A.3) with boundary conditions (A.8)–(A.9) for the potential (5.20)–(5.21). The values of W increase as the horizon position takes values closer and closer to the IR extremum at $\phi = v$. None of the black-hole solutions reach the local minimum of V at $\phi = -1$, they all bounce for $\phi > -1$. The position of the bounces and their progression as ϕ_h approaches v is clearer when looking at $W'(\phi)$, as is done in figures 30 and 31.

resulting superpotentials are displayed in figure 29, where only the branch above the bound is displayed and we set $v = 1$. The closer the horizon position ϕ_h is to $\phi = v$, the larger W and the closer the bounce is to $\phi = -1$. No finite-temperature solution reaches $\phi = -1$, no matter how close to $v = 1$ we set the horizon. This is further emphasised in figures 30 and 31 which display $W'(\phi)$ for different values of ϕ_h .

The conclusion is that all the black holes in this model miss the UV minimum at $v = -1$ and flow towards the intermediate maximum. The latter however violates the BF bound, so these states are not part of healthy theory. Apart from this, conceptually they are in the same class as the black holes with UV asymptotics at a maximum of V , which were discussed in the previous sections.

Acknowledgments

We thank Dio Anninos, Jerome Gauntlet, Jewel Kumar Ghosh, Lukas Witkowski, Aron Jansen, David Mateos, Ioannis Papadimitriou, Kyriakos Papadodimas, Christopher Rosen, Sergey Sibiryakov for discussion and comments.

This work was supported in part by the Advanced ERC grant SM-grav, No 669288.

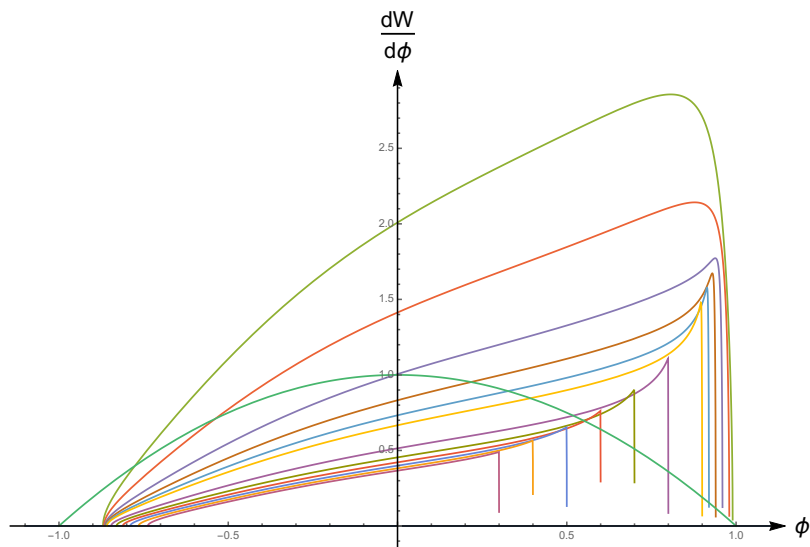


Figure 30. This figure shows the first derivative of the functions displayed in figure 30. The curve which extends from -1 to 1 corresponds to the zero-temperature solution. The remaining curves are the derivatives with respect to ϕ of the finite-temperature superpotentials displayed in figure 29. The rightmost zero (or the tendency of W' to vanish) corresponds to the horizon position which approaches $\phi = 1$ without reaching it and the leftmost zero corresponds to the bounce position. The closer the horizon is to $\phi = 1$ the closer the bounce gets to $\phi = -1$ but the superpotential never reaches -1 . This means that the zero-temperature minimum-to-minimum flow is not associated with a black hole. The bouncing points are better visualised in figure 31.

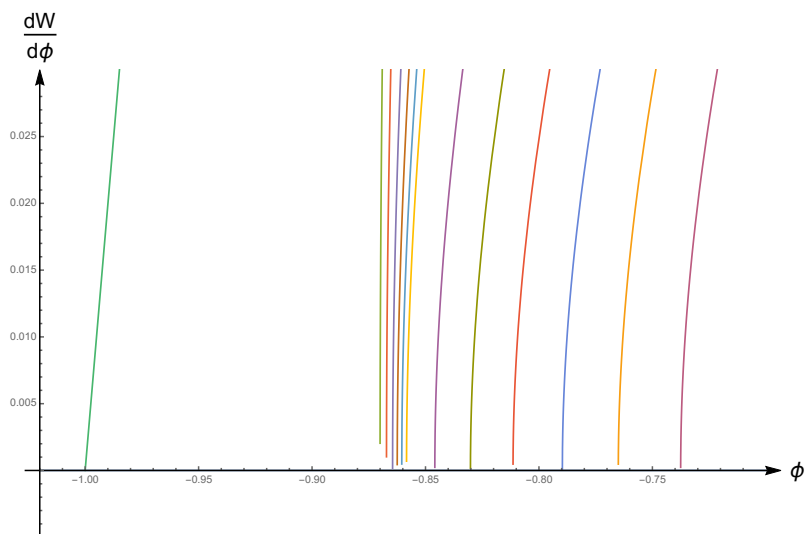


Figure 31. A closer look at the near-bounce region of figure 30. The leftmost curve corresponds to the zero temperature solution and the remaining curves are superpotentials at finite temperature. The derivative of each finite-temperature superpotential vanishes at a different bouncing point, where the each solution is glued to the next branch. The closer to -1 the zero of $dW/d\phi$ is, closer the horizon is to $\phi = 1$ and lower is the black-hole temperature. No finite-temperature solution reaches $\phi = -1$, thus showing that the minimum-to-minimum zero temperature flow from [12] does not have an associated black-hole solution with a running scalar.

A First order formalism

As for the zero-temperature solutions, we can introduce a function $W(\phi)$ such that, on-shell

$$\dot{A}(u) = -\frac{1}{2(d-1)}W(\phi(u)), \quad \dot{\phi}(u) = \frac{dW(\phi(u))}{d\phi} \quad (\text{A.1})$$

This can always be achieved piecewise in any region where $\phi(u)$ is monotonic, by inverting the relation $\phi(u) \rightarrow u(\phi)$. With the definition (A.1) equation (2.3d) is identically satisfied, and the remaining equations (2.3a)–(2.3c) become a system of scalar equations for the functions $W(\phi)$, $f(\phi) \equiv f(u(\phi))$ [31],

$$\left(\frac{1}{2} \left(\frac{dW}{d\phi} \right)^2 - \frac{d}{4(d-1)} W^2 \right) f + \frac{W}{2} \frac{df}{d\phi} \frac{dW}{d\phi} = V, \quad (\text{A.2})$$

$$\frac{d^2 f}{d\phi^2} \frac{dW}{d\phi} + \frac{df}{d\phi} \frac{d^2 W}{d\phi^2} - \frac{d}{2(d-1)} W \frac{df}{d\phi} = 0. \quad (\text{A.3})$$

For $f(\phi) = 1$, equation (A.2) becomes the usual superpotential equation.

Among the transformations (2.4) which leave the original equations of motion invariant, only (2.4c) affects (A.2) and (A.3). The corresponding transformations for $W(\phi)$ and $f(\phi)$ are

$$\left(\tilde{f}(\phi), \tilde{W}(\phi) \right) \rightarrow \left(f(\phi), W(\phi) \right) = \left(\lambda^2 \tilde{f}(\phi), \lambda^{-1} \tilde{W}(\phi) \right) \quad (\text{A.4})$$

and ϕ is left invariant.

A.1 The integration constants for the superpotential equations

The system (A.2)–(A.3) is third order¹⁴ and can be solved in terms of three integration constants which can be chosen in different ways. It is most convenient to view it as a boundary value problem for $f(\phi)$ and an initial value problem (at the black-hole horizon) for $W(\phi)$, and fix the integration constants by

$$f(0) = 1, \quad f(\phi_h) = 0, \quad W(\phi_h) = W_h. \quad (\text{A.5})$$

The first equation follows from the requirement that the leading asymptotic boundary metric is the same for all solutions. This leaves the two arbitrary parameters ϕ_h and W_h . As we will show below however, W_h is also fixed uniquely by the choice the horizon position ϕ_h , resulting in a one-parameter family of solutions parametrised by the value of ϕ_h .

A very simple way to see that the system (A.2)–(A.3) is third order and which is helpful to give boundary conditions for solving the system numerically goes as follows. By differentiating equation (A.2) with respect to ϕ , a term with $f''(\phi)$ is generated and can be used to eliminate $f''(\phi)$ from (A.3). The resulting equation is:

$$W'(\phi) \left[f(\phi) \left(\frac{dW(\phi)}{2-2d} + W''(\phi) \right) + f'(\phi) W'(\phi) \right] = V'(\phi), \quad (\text{A.6})$$

¹⁴Since both $\partial_\phi^2 W$ and $\partial_\phi^2 f$ appear, at first sight it may look as the system is fourth order, but this is not the case, as one can easily realise by algebraically solving equation (A.2) for $\partial_\phi W$ and inserting the result into equation (A.10).

Equations (A.2) and (A.6) demand three integration constants which completely specify $W(\phi)$ and $f(\phi)$. Evaluating (A.2) and (A.6) at the horizon we obtain

$$[Wf'W']_{\phi_h} = 2V(\phi_h), \quad (\text{A.7a})$$

$$[W'f'W']_{\phi_h} = V'(\phi_h). \quad (\text{A.7b})$$

For any choice of the integration constant $f'(\phi_h)$ and imposing that W is positive, the system (A.7) completely specifies $W(\phi_h)$ and $W'(\phi_h)$. Therefore one can choose boundary conditions close to the horizon as follows:

$$f(\phi_h + \delta\phi) = \delta\phi f'(\phi_h) + \mathcal{O}(\delta\phi^2) \quad \text{with} \quad f'(\phi_h)\delta\phi > 0, \quad (\text{A.8a})$$

$$W'(\phi_h + \delta\phi) = \sqrt{\frac{V'(\phi_h)}{f'(\phi_h)}} + \mathcal{O}(\delta\phi), \quad (\text{A.8b})$$

$$W(\phi_h + \delta\phi) = -\frac{2V(\phi_h)}{\sqrt{f'(\phi_h)V'(\phi_h)}} + \delta\phi \sqrt{\frac{V'(\phi_h)}{f'(\phi_h)}} + \mathcal{O}(\delta\phi^2). \quad (\text{A.8c})$$

The boundary conditions (A.8) will typically lead to $f(\phi_{\text{UV}}) \neq 1$, however this apparent inconvenience is immediately circumvented with a rescaling of the form (A.4) with parameter

$$\lambda = (f(\phi_{\text{UV}}))^{-1/2}. \quad (\text{A.9})$$

Another way to specify the integration constants is to explicitly identify the boundary theory source in $f(\phi)$ and the product TS in $W(\phi)$, making a more explicit contact with the thermodynamics. To proceed, we chose a value ϕ_h . First, notice that equation (A.3) can also be written equivalently as

$$\frac{d}{d\phi} \left[\frac{df}{d\phi} \frac{dW}{d\phi} e^{d\mathcal{A}(\phi)} \right] = 0 \quad (\text{A.10})$$

where

$$\mathcal{A}(\phi) = -\frac{1}{2(d-1)} \int^{\phi} d\hat{\phi} \frac{W}{W'} \quad (\text{A.11})$$

The function $\mathcal{A}(\phi)$ depends on an arbitrary additive constant but clearly this does not affect equation (A.10), and in fact it can be reabsorbed in one of the other integration constants. Integrating equation (A.10) gives

$$\frac{df}{d\phi} = -\frac{\tilde{D}}{W'} e^{-d\mathcal{A}(\phi)}, \quad (\text{A.12})$$

where \tilde{D} is a constant. As we mentioned above, an additive constant in $\mathcal{A}(\phi)$ can be absorbed in a redefinition of \tilde{D} and we can fix this redundancy (which does *not* correspond to one of the integration constants of the system (A.2)–(A.3)) so that $\mathcal{A}(\phi_h) = 0$:

$$\mathcal{A}(\phi) = -\frac{1}{2(d-1)} \int_{\phi_h}^{\phi} d\hat{\phi} \frac{W}{W'} \quad (\text{A.13})$$

Equation (A.12) can be further integrated to give

$$f(\phi) = \tilde{F} - \tilde{D} \int_0^\phi \frac{dy}{W'(y)} e^{-dA(y)}. \quad (\text{A.14})$$

As we will see in a moment, the second term vanishes as $\phi \rightarrow 0$. Therefore, imposing the boundary conditions for $f(\phi)$ from equation (A.5) fixes

$$\tilde{F} = 1, \quad \tilde{D} = \left(\int_0^{\phi_h} \frac{dy}{W'(y)} e^{-dA(y)} \right)^{-1}. \quad (\text{A.15})$$

Finally, evaluating equation (A.2) at ϕ_h and using (A.12) we find

$$W(\phi_h) = -\frac{2V(\phi_h)}{\tilde{D}}. \quad (\text{A.16})$$

This condition is equivalent to the requirement that the horizon is regular, and it fixes the remaining free parameter W_h in the initial conditions (A.5).

To summarise we have found that, for regular asymptotically AdS black holes, solutions to the system (A.2)–(A.10) depend on a single continuous free parameter, which we can be taken to be the horizon position ϕ_h . In the next subsection we will clarify the physical meaning of the quantity $\tilde{D}(\phi_h)$.

A.2 Lower bounds on the superpotential

At zero temperature, an important property of $W(\phi)$ is that it is bounded from below. This can be shown by considering equation (A.2) with $f(\phi) \equiv 1$ and algebraically solving for W :

$$|W(\phi)| = \sqrt{\frac{4(d-1)}{d} \left(\frac{1}{2}W'^2 - V(\phi) \right)} \geq \sqrt{-\frac{4(d-1)}{d}V(\phi)} \equiv B(\phi) > 0. \quad (\text{A.17})$$

When $V(\phi)$ is negative, W can never approach zero, meaning that the branches with $W > 0$ and $W < 0$ which solve (A.17) are disconnected. Furthermore, they are physically equivalent as one can pass from one to the other by the transformation $(u, W) \rightarrow (-u, -W)$. Therefore, without loss of generality we can choose a positive W , meaning that $\dot{A}(u)$ is negative, and see (A.17) as a lower bound on W and not just on its absolute value:

$$W(\phi) \geq B(\phi) \equiv \sqrt{-\frac{4(d-1)}{d}V(\phi)} > 0. \quad (\text{A.18})$$

Another useful property to understand the numerical results shown in figures 29, 30 and 31 which is related to the normalisation of $f(\phi)$ is that W can lie below $B(\phi)$ defined in (A.18). It is instructive to algebraically solve equation (A.2) for $W(\phi)$ keeping into account the positivity of W which is necessary for \dot{A} to be negative

$$W(\phi) = \frac{(d-1)}{d} \frac{f'W'}{f} + \sqrt{\left(\frac{(d-1)}{d} \frac{f'W'}{f} \right)^2 + \frac{2(d-1)}{d} \left(W'^2 - \frac{2V}{f} \right)} \quad (\text{A.19})$$

As long as $V(\phi)$ is negative, it is guaranteed that W is real at stationary points, i.e. when W' vanishes. Negativity of $V(\phi)$ also guarantees that giving W a positive value at any fixed $\phi = \phi_0$ implies the positivity of W for all ϕ . In contrast to the zero temperature case, the function $f(\phi)$ permits the existence of solutions such that the limit

$$\left(\frac{f'W'}{f}\right)^2 \gg \frac{2d}{d-1} \left(W'^2 - \frac{2V}{f}\right) \tag{A.20}$$

holds, thus including arbitrarily small superpotentials in the space of solutions. On the other hand, from (A.19) it follows that the value of $W(\phi)$ at critical points, i.e. those where $W'(\phi)$ vanishes, is set by $f(\phi)$ and $V(\phi)$,

$$W' = 0 \quad \text{and} \quad \frac{f'W'}{f} = 0 \quad \implies \quad W = \sqrt{-\frac{4(d-1)V}{d} \frac{1}{f}} \tag{A.21}$$

one of the few remaining properties from the zero temperature setup.

A.3 Near-boundary solution: universal part

To obtain the solution close to the boundary, we solve the system of equations (A.2)–(A.10) perturbatively close to $\phi = 0$, where $f \simeq 1$. To lowest order we obtain the analytic part W_0 of the superpotential,

$$W_0^\pm(\phi) = \frac{2(d-1)}{\ell} + \frac{\Delta_\pm}{2\ell} \phi^2 + O(\phi^4), \tag{A.22}$$

which coincide with the perturbative expansion in the vacuum solution, see e.g. [12, 33, 34]. The two choices, Δ_- or Δ_+ , in the quadratic term, correspond respectively to a flow driven by a source or by a vev of the operator dual to ϕ . Generically we will be describing flows with non-zero source, therefore from now on we will choose $W_0 = W_0^-$. Vev-driven flows will play an important role however, and will be discussed separately later on.

Up to higher order corrections, we can replace W by W_0 in \mathcal{A} in (A.12) and close to the boundary we find

$$\mathcal{A}(\phi) = \mathcal{A}_0 - \frac{1}{\Delta_-} \log |\phi| + O(\phi) \tag{A.23}$$

where \mathcal{A}_0 is a constant which is also completely determined by equation (A.13) and depends only on ϕ_h .

Using equations (A.22) and (A.23) we can approximate equation (A.14) for close to $\phi = 0$,

$$f(\phi) = 1 - \frac{\tilde{D} e^{-d\mathcal{A}_0} \ell}{d} |\phi|^{\frac{d}{\Delta_-}}. \tag{A.24}$$

As announced, the ϕ -dependent part vanishes as $\phi \rightarrow 0$, and f satisfies the correct boundary condition.

Finally, integrating equations (A.1) give the expected leading order scaling,

$$A(u) = -\frac{u}{\ell} + O(e^u), \quad \phi(u) = j \ell^{\Delta_-} e^{\Delta_- u/\ell} + \dots, \quad u \rightarrow -\infty, \tag{A.25}$$

and introduces as additional free parameter j , the constant corresponding to the source of the dual operator, whose choice is part of the boundary data one has to give to fix the holographic theory. The other piece of boundary data is the leading asymptotic behaviour of the boundary metric: in writing equation (A.25) we have chosen the constant term in A to vanish, so that in the holographic dictionary the dual field theory metric is $\eta_{\mu\nu}$ with unit coefficient.

Thus, the full solution $(A(u), \phi(u), f(u))$ depends on only two parameters: the value of the source j and the horizon position ϕ_h . The black-hole temperature and entropy density are therefore determined by these two quantities. In the next section we will make this dependence more explicit.

A.4 Dimensionless temperature and entropy

Equations (A.24)–(A.25) reproduce the expected scaling (2.11) near the boundary. Comparing equations (A.24) and (2.11c), and using the relation (A.12) we can relate \tilde{D} to the black-hole temperature T and entropy density s :

$$\tilde{D}e^{-dA_0} = \frac{1}{M_p^{d-1}} \frac{Ts}{|j|^{d/\Delta_-}}. \tag{A.26}$$

As the right hand side depends only on ϕ_h , the same must be true for the combination on the right hand side. In fact we will now show that the two quantities

$$\mathcal{T} \equiv \frac{T}{|j|^{1/\Delta_-}}, \quad \mathcal{S} \equiv \frac{s}{|j|^{(d-1)/\Delta_-}} \tag{A.27}$$

are function of ϕ_h only. They represent the temperature and entropy density in units of the UV source j .¹⁵

To see that the \mathcal{T} and \mathcal{S} defined above depend on ϕ_h only, and not on both ϕ_h and j , we use the invariance property (2.4a)–(2.4b),

$$u \rightarrow u' = u + v, \quad A(u) \rightarrow A'(u') = A(u') + \bar{A}. \tag{A.28}$$

Starting from a solution, these transformation produces a new solution of the same form (2.2) but with different UV data, temperature and entropy: the source term transforms as

$$j \rightarrow e^{\Delta_- v} j \tag{A.29}$$

and the horizon data as

$$u_h \rightarrow u_h + v, \quad e^{A(u_h)} \rightarrow e^{A(u_h+v)+\bar{A}}, \quad \dot{f}(u_h) \rightarrow \dot{f}(u_h + v) \tag{A.30}$$

However if we want to keep the boundary metric fixed to be $\eta_{\mu\nu}$ we have to choose $\bar{A} = v/\ell$, as we have to cancel a constant term in the asymptotic expansion of $A(u)$, equation (A.25). Taking this into account, we find that under (A.28)

$$T \rightarrow e^v T, \quad s \rightarrow e^{(d-1)v} s \tag{A.31}$$

¹⁵They are the finite-temperature analog of the dimensionless curvature parameter \mathcal{R} used in [16].

On the other hand, the horizon value ϕ_h is unaffected by this transformation, so all solutions related by the transformation (A.28) have the same ϕ_h and the same values of the combinations (A.27). We conclude that both \mathcal{T} and \mathcal{S} are determined solely by the horizon field value ϕ_h . Since \mathcal{T} (as opposed to ϕ_h) is directly related to the horizon data (T, j) , it is useful to invert the relation $\mathcal{T}(\phi_h)$ (at least piecewise) and use \mathcal{T} to parametrise black-hole solutions with the same ϕ_h .

A.5 Near-boundary solution: sub-leading term

The sub-leading, non-analytic contributions to $W(\phi)$ beyond the leading behaviour (A.22) contains the information about the vevs, see [12, 34] and gives finite contributions to the free energy. It can be obtained by setting $W = W_0 + \delta W$, linearizing equation (A.2) around the solution (A.22) and keeping in mind that $f = 1 + O(\phi^{d\Delta_-})$. The resulting equation is

$$W_0' \delta W' - \frac{d}{2(d-1)} W_0 \delta W_0 = (f_0 - 1)V + f_0' W_0' W_0 \quad (\text{A.32})$$

Using the expression for f_0 in equation (A.24) and evaluating V at the origin (recall $V(0) = -d(d-1)\ell^{-2}$) we find that the right hand side vanishes at the lowest order ϕ^{d/Δ_-} , and it starts at order ϕ^{d/Δ_-+2} . To lowest order the solution for δW is

$$\delta W = \ell^{-1} C(\phi_h) |\phi|^{d/\Delta_-} (1 + O(\phi^2)) \quad (\text{A.33})$$

where $C(\phi_h)$ is an integration constant which, following the discussion in section A.1, is determined by the horizon value ϕ_h . Equation (A.33) has exactly the same form as the generic non-analytic part of the superpotential at zero temperature, the only difference being that the integration constant is fixed by the horizon position in field space.

Rather than ϕ_h , we will find it convenient to consider $\mathcal{T}(\phi_h)$ as the independent parameter, since it directly relates to the boundary quantities T, j . Therefore we will often write $C(\mathcal{T})$ instead of $C(\phi_h)$.

In terms of the domain-wall coordinate u , using the leading order behaviour $\phi \simeq j e^{\Delta_- u}$, we find that δW scales as e^{du} . More precisely,

$$\delta W \simeq \ell^{-1} C(\mathcal{T}) |j|^{\Delta_+/\Delta_-} e^{du/\ell}, \quad u \rightarrow -\infty. \quad (\text{A.34})$$

Finally, we can integrate the first order equation $\dot{\phi}(u) = W'(\phi(u))$ to sub-leading order as $u \rightarrow -\infty$, to find

$$\phi(u) = j \ell^{\Delta_-} e^{\Delta_- u} + \dots + \phi_+ \ell^{\Delta_+} e^{\Delta_+ u}, \quad \phi_+ = \frac{d}{\Delta_-} \frac{C(\mathcal{T})}{(d - 2\Delta_-)} |j|^{\Delta_+/\Delta_-} \text{sign}(j). \quad (\text{A.35})$$

Therefore, $C(\mathcal{T})$ plays the role as a dimensionless vev parameter, $\langle O \rangle / |j|^{\Delta_+/\Delta_-}$, which depends on \mathcal{T} (or equivalently, on ϕ_h) only.

A.6 Superpotentials for vev-driven flows

To conclude this appendix we discuss solutions with $j = 0$ which nevertheless describe a non-trivial flow. For these solutions, the appropriate (analytic part of the) superpotential

corresponds to the choice Δ_+ in equation (A.22),

$$W^+ = \frac{2(d-1)}{\ell} + \frac{\Delta_+}{2\ell} \phi^2 + O(\phi^4) \tag{A.36}$$

In this case, a similar analysis to the one in section A.5 shows that a sub-leading, non-analytic part is *not* allowed (see [12, 34] for a more extended discussion). Therefore the solution of the kind W^+ is unique, it is completely fixed by its analytic expansion, and does not admit a continuous parameter family of deformations.

Since they have, $j = 0$ these solutions correspond to taking the limit $\mathcal{T} \rightarrow +\infty$, which occurs (at most) at isolated special values of ϕ_h .

After integrating the first order flow equations for $A(u)$ and $\phi(u)$, the full solution is parametrised by the vev ϕ_+ (which now is a free integration constant), and the near-boundary expansion of $\phi(u)$, solving $\dot{\phi} = W'$, is

$$\phi(u) = \phi_+ \ell^{\Delta_+} e^{\Delta_+ u} + \dots, \quad \langle O \rangle = (M_P \ell)^{d-1} (d - 2\Delta_-) \phi_+. \tag{A.37}$$

Fixing ϕ_+ also fixes the temperature of the solution since, by a similar argument as the one presented in section A.1, it is easy to show that for all such solutions¹⁶) the ratio $T/\phi_+^{1/\Delta_+}$ is fixed.

B The on-shell action

The on-shell Euclidean action is finite if we include the appropriate counter-term action in equation (2.1). It was systematically analysed in [31]. Here we discuss the three terms in (2.1) separately. In doing so, each term is finite only if we introduce a UV cut-off $u = u_{UV}$ in the asymptotic boundary region. In the final step one removes the cut-off by taking $u_{UV} \rightarrow -\infty$.

Bulk action. This can be written in terms of the UV asymptotic data, as follows. The trace of the Einstein's equations of motion following from (2.1) lead to the following expression for the Ricci scalar:

$$R = \frac{\dot{\phi}^2}{2f} + \frac{d+1}{d-1} V, \tag{B.1}$$

which can be used to eliminate R from the bulk part of the action S^{bulk} , i.e. the first term in equation (2.1), leaving

$$S^{\text{bulk}} = -\frac{2M_P^{d-1}}{d-1} \int_{\mathcal{M}} d^{d+1}x e^{dA(u)} V(\phi(u)) \tag{B.2}$$

Using equation (2.3b) to eliminate V and equation (2.3d) to eliminate $\dot{\phi}^2$ we find

$$\begin{aligned} S^{\text{bulk}} &= 2M_P^{d-1} \int_{u_{UV}}^{u_h} \int d^d x e^{dA} \left[\dot{f} \dot{A} + (d\dot{A}^2 + \ddot{A}) f \right] \\ &= 2M_P^{d-1} \beta V_{d-1} \left[e^{dA} \dot{A} f \right]_{u_{UV}}^{u_h} \end{aligned} \tag{B.3}$$

¹⁶That is, all solutions with the same value of ϕ_h corresponding to one of the special points in parameter space.

where V_{d-1} is the $d - 1$ -dimensional spatial volume (which we take to be finite, e.g. by considering the system in a spatial square box of size L). Because $f(u_h) = 0$, the horizon does not contribute, and the contribution from the bulk part of the action is:

$$S^{\text{bulk}} = -2M_P^{d-1}\beta V_{d-1} \left[e^{dA} \dot{A}f \right]_{u=u_{\text{UV}}} \quad (\text{B.4})$$

Gibbons-Hawking-York term. To compute the Gibbons-Hawking-York boundary term S_{GHY} from (2.1) we define the outward-oriented unit normal vectors to the boundary at $u = u_{\text{UV}}$. Notice that we should *not* include a GHY term at the horizon, since this is a regular point in the interior of the Euclidean geometry, rather than a boundary.

The unit normal to the boundary is given by

$$n^a = -\sqrt{f} \delta_u^c \quad \text{at } u = u_{\text{UV}} \quad (\text{B.5})$$

and the extrinsic curvature is

$$K = \nabla_a n^a = -\frac{1}{\sqrt{f}} \left(\dot{f} + d\dot{A}f \right) \quad (\text{B.6})$$

The on-shell expression for S_{GHY} from (2.1) for solutions of the form (2.2) is then

$$S_{\text{GHY}}^{\text{on-shell}} = M_P^{d-1}\beta V_{d-1} \left[e^{dA} \left(\dot{f} + 2d\dot{A}f \right) \right]_{u=u_{\text{UV}}} . \quad (\text{B.7})$$

Combining the two terms (B.4) and (B.7) we obtain the regularised on-shell action,

$$S_{\text{on-shell}}^{\text{reg}} = M_P^{d-1}\beta V_{d-1} \left[e^{dA} \left(\dot{f} + 2(d-1)\dot{A}f \right) \right]_{u=u_{\text{UV}}} . \quad (\text{B.8})$$

This expression is divergent as we let $u_{\text{UV}} \rightarrow -\infty$. Before we remove the cut-off we must add appropriate counter-terms, which we discuss below.

Counter-term action. The counter-terms are universal, and are written in the general form [34, 35]

$$S_{\text{ct}} = M_P^{d-1} \int_{u=u_{\text{UV}}} d^d x \sqrt{\gamma} \left[W_{\text{ct}}(\phi) + U_{\text{ct}}(\phi) R^{(\gamma)} \right] \quad (\text{B.9})$$

where u_{UV} is a UV cut-off, γ is the induced metric on $u = u_{\text{UV}}$, $R^{(\gamma)}$ its intrinsic curvature, and W and U are appropriate functions of the scalar field.

In the case at hand, the $u = \text{const.}$ hyper-surfaces are flat, so $R^{(\gamma)} = 0$ and the curvature counter-term vanishes identically. The appropriate function $W_{\text{ct}}(\phi)$ must solve the *superpotential equation*,

$$\frac{1}{2} \left(\frac{dW_{\text{ct}}}{d\phi} \right)^2 - \frac{d}{4(d-1)} W_{\text{ct}} = V \quad (\text{B.10})$$

In the case $d/2 > \Delta_- > d/4$, to which we restrict here, the solution for W close to the boundary has the universal form:

$$W_{\text{ct}}(\phi) = W_0(\phi) + \frac{C_{\text{ct}}}{\ell} |\phi|^{\frac{d}{\Delta_-}} + \dots, \quad (\text{B.11})$$

up to term which vanish faster than e^{-dA} as $\phi \rightarrow 0$. In the expression above, W_0 is the analytic solution whose expression can be found in equation (A.22), taken up to quadratic order; C_{ct} is an arbitrary constant which encodes the scheme dependence of the subtraction. We choose to work in a minimal scheme, in which we set $C_{\text{ct}} = 0$. The resulting counter-term is quadratic in ϕ and we will denote it by

Evaluating the counter-term on-shell gives:

$$S_{\text{ct}} = M_p^{d-1} \beta V_{d-1} \left[e^{dA} \sqrt{f} W_0 \right]_{u=u_{\text{UV}}} . \tag{B.12}$$

B.1 Calculation of the free energy

The renormalised Free energy is found by adding the counter-term (B.12) to the regularised on-shell action (B.8),

$$\beta \mathcal{F} = M_P^{d-1} \beta V_{d-1} \left[e^{dA} \left(\dot{f} + 2(d-1) \dot{A} f + W_0 \sqrt{f} \right) \right]_{u=u_{\text{UV}}} \tag{B.13}$$

Let us look at the two terms in the square brackets separately.

1. Using equations (2.8) and (2.9) we can rewrite the first term simply as follows:

$$M_P^{d-1} \beta V_{d-1} \left[e^{dA} \dot{f} \right]_{u_{\text{UV}}} = -\beta T s V_{d-1}. \tag{B.14}$$

2. The second term can be recast in a simpler form if we revert to the first order formalism developed in appendix A, and we write \dot{A} in terms of the finite temperature superpotential W using equation (A.1):

$$\left(2(d-1) \dot{A} f + \sqrt{f} W_0 \right) = -\sqrt{f} \left(W \sqrt{f} - W_0 \right) e^{dA}. \tag{B.15}$$

As shown in appendix A, close to the boundary the superpotential takes the form

$$W = W_0(\phi) + \frac{C(\mathcal{T})}{\ell} |\phi|^{d/\Delta_-} + \dots \tag{B.16}$$

up to terms which vanish faster than $e^{du/\ell}$ as $u \rightarrow -\infty$. In the equation above, $W_0(\phi)$ is the same universal power-series which solves the zero-temperature superpotential equation (B.10), and which enters the counter-term (B.11).

We now expand \sqrt{f} and e^{dA} using (2.11),

$$\sqrt{f} \simeq 1 - \frac{\ell D}{2d} e^{du/\ell} + O\left(e^{2du/\ell}\right), \quad e^{dA} \simeq e^{-du} (1 + O(e^{2\Delta_- u/\ell})). \tag{B.17}$$

Inserting the expansions (B.16) and (B.17), the right hand side of equation (B.15) takes the form,

$$\sqrt{f} \left(W \sqrt{f} - W_0 \right) e^{dA} \simeq -W_0 \frac{\ell D}{2d} - \frac{C(\mathcal{T})}{\ell} |\phi|^{d/\Delta_-} e^{dA} \rightarrow \frac{(d-1)}{d} \frac{T s}{M_p^{d-1}} - C(\mathcal{T}) \ell^{d-1} |j|^{d/\Delta_-} \tag{B.18}$$

where in the last step we have taken the limit $u_{\text{UV}} \rightarrow -\infty$ and we have expressed D in terms of $T s$ by equation (2.9).

Putting everything together we finally arrive at the expression for the free energy,

$$\mathcal{F} = -\frac{T s}{d} V_{d-1} - (M_p \ell)^{d-1} C(\mathcal{T}) |j|^{d/\Delta_-} V_{d-1}. \tag{B.19}$$

B.2 Thermal vev

In this appendix we show the validity of the relation

$$\frac{\partial \mathcal{F}(T, j)}{\partial j} = -\langle O \rangle V_{d-1} \quad (\text{B.20})$$

where the dual operator vev $\langle O \rangle$ is related to $C(\mathcal{T})$ and j by equation (A.35),

$$\langle O \rangle = (M_p \ell)^{d-1} \frac{d}{\Delta_-} C(\mathcal{T}) |j|^{\Delta_+/\Delta_-} \text{sign}(j). \quad (\text{B.21})$$

For simplicity of notation we suppose $j > 0$, but the final result holds for either sign of j .

We start from the expression (2.28),

$$\mathcal{F} = T^d \left(-\frac{\sigma(\mathcal{T})}{d} + \gamma(\mathcal{T}) \right), \quad \gamma(\mathcal{T}) \equiv (M_p \ell)^{d-1} V_{d-1} \frac{C(\mathcal{T})}{\mathcal{T}^d}, \quad \mathcal{T} = \frac{T}{j^{\Delta_-}}. \quad (\text{B.22})$$

First, we differentiate \mathcal{F} with respect to T , and use the identities

$$\frac{\partial \mathcal{T}}{\partial T} = \frac{1}{j^{\Delta_-}}, \quad \frac{\partial \mathcal{T}}{\partial j} = -\frac{1}{\Delta_- j} \mathcal{T}, \quad (\text{B.23})$$

to obtain

$$\frac{\partial \mathcal{F}}{\partial T} = -T^{d-1} \left[\left(\frac{\sigma'}{d} + \gamma' \right) \mathcal{T} + d \left(\frac{\sigma}{d} + \gamma \right) \right], \quad (\text{B.24})$$

where a prime denotes a derivative with respect to the argument, \mathcal{T} . On the other hand, we make use of the thermodynamic relation

$$s = -\frac{1}{V_{d-1}} \frac{\partial \mathcal{F}}{\partial T}, \quad (\text{B.25})$$

to replace the left hand side of equation (B.24) by $-T^{d-1} \sigma$. This leads to the relation

$$\left(\frac{\sigma'}{d} + \gamma' \right) = -d\gamma. \quad (\text{B.26})$$

Next, we differentiate equation (B.22) with respect to j ,

$$\frac{\partial \mathcal{F}(T, j)}{\partial j} = T^d \left(\frac{\sigma'}{d} + \gamma' \right) \frac{\mathcal{T}}{\Delta_- j} = -\frac{d}{\Delta_-} \frac{T^d}{j} \gamma, \quad (\text{B.27})$$

where we have used the result (B.26) in the second equality. Finally, using the definition of γ from equation (B.22), we obtain,

$$\frac{\partial \mathcal{F}(T, j)}{\partial j} = -(M_p \ell)^{d-1} V_{d-1} \frac{d}{\Delta_-} C(\mathcal{T}) j^{\Delta_+/\Delta_-} \quad (\text{B.28})$$

Comparing the right hand side with the equation above with the relation (B.21), we obtain the desired result (B.20).

C Determination of the free energy using scalar variables

C.1 The phase variables

We define the two phase variables X and Y as:

$$X(\phi) \equiv \frac{\gamma \phi'}{d A'}, \quad Y(\phi) \equiv \frac{1}{d} \frac{g'}{A'}, \quad (\text{C.1})$$

where the function g is defined as $g = \log f$ and the constant γ is given by,

$$\gamma = \sqrt{\frac{d}{2(d-1)}}. \quad (\text{C.2})$$

These functions satisfy:

$$\frac{dX}{d\phi} = -\gamma (1 - X^2 + Y) \left(1 + \frac{1}{2\gamma} \frac{1}{X} \frac{d \log V}{d\phi} \right), \quad (\text{C.3})$$

$$\frac{dY}{d\phi} = -\gamma (1 - X^2 + Y) \frac{Y}{X}. \quad (\text{C.4})$$

This second order system is sufficient to determine all of the thermodynamic properties (and dissipation) of the gravitational theory [27]. This is a reduction of the fifth order Einstein-scalar system to an equivalent second order system.

It is straightforward to show that these equations combined with the following three,

$$\frac{dA}{du} = -\frac{1}{\ell} e^{-\gamma \int_0^\phi X(t) dt}, \quad (\text{C.5})$$

$$\frac{d\phi}{du} = -\frac{1}{\ell} \frac{d}{d\phi} X(\phi) e^{-\gamma \int_0^\phi X(t) dt}, \quad (\text{C.6})$$

$$\frac{dg}{du} = -\frac{1}{\ell} \frac{d}{d\phi} Y(\phi) e^{-\gamma \int_0^\phi X(t) dt}, \quad (\text{C.7})$$

solve the original Einstein equations in the domain-wall variables defined by the ansatz (2.2). The solution in the conformal coordinates is found by the change of variables $du = \exp(A) dr$,

$$ds^2 = e^{2A(r)} (f^{-1}(r) dr^2 + dx_{d-1}^2 + dt^2 f(r)), \quad \Phi = \Phi(u). \quad (\text{C.8})$$

One can also express g and A in terms of the phase variables directly from the definitions (C.1):

$$A(\phi) = A(\phi_c) + \frac{\gamma}{d} \int_{\phi_c}^\phi \frac{d\tilde{\phi}}{X}, \quad (\text{C.9})$$

$$f(\phi) = \exp \left(\gamma \int_0^\phi \frac{Y}{X} d\tilde{\phi} \right). \quad (\text{C.10})$$

Here Φ_c denotes a surface near the boundary where we will apply the UV matching conditions of the TG and the BH solution in the following. Another useful equation relates the scalar potential to the phase variables, that follows from (C.3):

$$V(\phi) = \frac{d(d-1)}{\ell^2} (1 + Y - X^2) e^{-2\gamma \int_0^\phi (X(t) - \frac{Y(t)}{2X(t)}) dt}. \quad (\text{C.11})$$

The precise form of the overall coefficient follows from inserting (C.5), (C.6) and (C.7) in the Einstein's equations.

The temperature T and the entropy density s of the black-hole are given by

$$T(\phi_h) = \frac{\ell}{4\pi(d-1)} e^{A(\phi_h)} V(\phi_h) e^{\gamma \int_0^{\phi_h} X(\phi) d\phi}, \quad (\text{C.12})$$

$$s = 4\pi M_p^{d-1} e^{(d-1)A(\phi_h)}. \quad (\text{C.13})$$

In the first equation we used

$$-4\pi T = f'(r_h) = \left. \frac{df}{d\phi} \frac{d\phi}{du} \frac{du}{dr} \right|_{\phi_h} = \frac{d}{\ell} Y(\phi_h) e^{A(\phi_h) + \gamma \int_0^{\phi_h} (\frac{Y}{X} - X) d\bar{\phi}}, \quad (\text{C.14})$$

and the fact that $Y(\phi)$ diverges as $\phi \rightarrow \phi_h$ to express it in terms of the same limit of the potential (C.11). As $A(\phi)$ from (C.9) diverges as the cut-off ϕ_c is removed, equations that explicitly involve $A(\phi_h)$ are not very efficient practically. Instead, one can derive the following equations with no reference to $A(\phi)$ after a little bit of algebra:

$$T = T_{\text{ref}} \left(V(\phi_h) \ell^2 e^{\gamma \int_0^{\phi_h} X} \right)^{1-\frac{1}{d}} Y_0(\phi_h)^{\frac{1}{d}}, \quad s = s_{\text{ref}} \left(Y_0(\phi_h)^{-1} V(\phi_h) \ell^2 e^{\gamma \int_0^{\phi_h} X} \right)^{\frac{1}{d}-1}, \quad (\text{C.15})$$

where T_{ref} and s_{ref} are constants given by:

$$T_{\text{ref}} = \frac{d^{\frac{1}{d}} (d-1)^{\frac{d}{d-1}} |j|^{\frac{1}{\Delta_-}}}{4\pi}, \quad s_{\text{ref}} = 4\pi (d(d-1))^{1-\frac{1}{d}} (M_P \ell)^{d-1} |j|^{\frac{d-1}{\Delta_-}}. \quad (\text{C.16})$$

Finally we note that the variable X is related to the superpotential defined in section A as

$$W(\phi) = \frac{2(d-1)}{\ell} e^{-\gamma \int_0^{\phi} X(t) dt}. \quad (\text{C.17})$$

C.2 UV and IR asymptotics

We first discuss the UV asymptotics in the zero T solution. Equation (C.3) near $\phi = 0$ yields

$$X_0(\phi) = -X_{\pm} \phi + \mathcal{O}(\phi^2) \quad (\text{C.18})$$

where

$$X_{\pm} = \frac{\gamma}{2} \left(1 \pm \sqrt{1 + \frac{4m^2 \ell^2}{d^2}} \right). \quad (\text{C.19})$$

Here the integration constant of equation (C.3) is not visible in the Taylor expansion near the UV. In fact it is given by a non-analytic term. As discussed in [27] this integration constant is completely fixed by the choice of asymptotics in the IR. In passing we note the relationship:

$$\frac{\gamma}{X_{\pm}} = \frac{d}{\Delta_{\pm}}. \quad (\text{C.20})$$

As mentioned above the relevant deformations correspond to $m^2 < 0$. The BF bound can be read off from (C.19) as,

$$\frac{m^2 \ell^2}{d^2 \xi} \geq -\frac{1}{2}. \quad (\text{C.21})$$

By computing the expansion of ϕ near the boundary one learns that the choice $X \rightarrow -X_- \phi$ corresponds to a deformation of the UV conformal theory by a source, and $X \rightarrow -X_+ \phi$ corresponds to a VeV, hence spontaneous breaking of conformal symmetry. We will assume deformation by a source below.

At finite temperature, solution of (C.3) and (C.4) near ϕ yields,

$$Y(\phi) = Y_0(\phi_h) \phi^{\frac{d}{\Delta_-}} + \dots \tag{C.22}$$

$$X(\phi) = X_0(\phi) + \delta X_0(\phi_h) \phi^{\frac{d}{\Delta_-} - 1} + \dots \tag{C.23}$$

Here, δX is defined as the deformation due to the presence of the BH, i.e. $\delta X = 0$ for the TG solution. We indicated the dependence of the integration constants in the UV on the location of the horizon ϕ_h .

Using these asymptotics in (C.5) and (C.6) we find

$$\phi(u) = \phi_- e^{\Delta_- \frac{u}{\ell}} + \dots \tag{C.24}$$

$$A(u) = \tilde{A}_0 - \frac{u}{\ell} + \dots \tag{C.25}$$

where $u \rightarrow -\infty$ corresponds to the boundary and ϕ_0 and \tilde{A}_0 are the integration constants. Because the equations of motion are invariant under the change of variables $u \rightarrow u + const$, the integration constant \tilde{A}_0 can be set to zero. This is what we will do in the following. Changing to the conformal frame near the boundary where $r/\ell = \exp(u/\ell)$, we have

$$\phi(r) = jr^{\Delta_-} + \phi_+ r^{\Delta_+} + \dots \tag{C.26}$$

$$A(r) = -\log \frac{r}{\ell} + \dots \tag{C.27}$$

We further note the constant limit

$$\lim_{\phi \rightarrow 0} e^{A(\phi)} \phi^{\frac{\gamma}{\Delta_-}} = \ell |j|^{\frac{1}{\Delta_-}}. \tag{C.28}$$

UV asymptotics of X and Y cannot depend on any other integration constant by the following simple argument. One solves (C.3) and (C.4) starting from the horizon. A priori one expects two integration constants. Location of the horizon can be viewed as one, therefore there remains one. However, demanding a regular horizon of the form

$$f(\phi) = \text{const.} \times (\phi_h - \phi), \quad \text{near } \phi_h, \tag{C.29}$$

from the equations (C.10) and (C.4) means that this remaining integration constant is completely fixed, as one should require,

$$X(\phi) = -\frac{1}{2\gamma} \frac{V'(\phi_h)}{V(\phi_h)} + \mathcal{O}(\phi_h - \phi); \quad Y(\phi) = -\frac{X(\phi_h)}{\gamma(\phi_h - \phi)} + \mathcal{O}(1), \tag{C.30}$$

near $\phi \approx \phi_h$. In general, this argument shows that one does not have to worry about the other integration constant which can be thought of as the source j — that we discuss below — if one derives the thermodynamics directly from the X, Y system. Below, we show this in more detail.

C.3 The free energy

One can calculate the free energy directly from the on-shell value of the GR action by using the solution expressed in terms of the phase variables above. The blackhole metric in the ϕ frame reads,

$$ds_{\text{BH}}^2 = B^2(\phi) \left(dt^2 F(\phi) + d\vec{x}^2 + \frac{d\phi^2}{F(\phi)D(\phi)^2} \right). \quad (\text{C.31})$$

Here the various metric functions are defined as follows:

$$B(\phi) = B(\phi_0) e^{\frac{\gamma}{d} \int_{\phi_c}^{\phi} \frac{d\tilde{\phi}}{X}}, \quad (\text{C.32})$$

$$D(\phi) = -\frac{d}{\gamma \ell} X(\phi) B(\phi) e^{-\gamma \int_0^{\phi} d\tilde{\phi} X}, \quad (\text{C.33})$$

$$F(\phi) = e^{\gamma \int_0^{\phi} d\tilde{\phi} \frac{Y}{X}}, \quad (\text{C.34})$$

where ϕ_c is a UV cut-off that we will remove at the end of the calculation. They are obtained directly from the expressions for the metric functions defined in the text in terms of the radial variable r , viz. (C.9), (C.10) and (C.5)–(C.7). We call the metric functions in λ with the capital letters to distinguish them from the analogous functions of r . The relations are explicitly given by $B(\phi) = b(r(\phi))$, $F(\phi) = f(r(\phi))$ where r is determined by

$$r(\phi) = \int_0^{\phi} \frac{d\tilde{\phi}}{D(\tilde{\phi})}.$$

The expressions above completely determine the map between the r -frame and the ϕ -frame.

Einstein contribution. We first compute the Einstein (bulk) contribution to the free energy, i.e. the first term in equation (2.1). The bulk on-shell action S^{bulk} (after using the Einstein's equations) is generally given by the frame-independent expression,

$$S^{\text{bulk}} = \frac{2}{d-1} M^{d-1} \int_{\mathcal{M}} \sqrt{g} V. \quad (\text{C.35})$$

\mathcal{M} is the manifold with a boundary. We regulate the integral in the ϕ -frame by placing a cut-off at ϕ_c . Thus, using the metric functions defined above, one obtains the following expression in the ϕ variable

$$S^{\text{bulk}} = \frac{2}{d-1} M^{d-1} \beta V_{d-1} \int_{\phi_c}^{\phi_h} B(\phi)^{d+1} V(\phi) D(\phi)^{-1}. \quad (\text{C.36})$$

We now substitute the expression for $D(\phi)$, $B(\phi)$ and $V(\phi)$ from (C.33), (C.32) and (C.11), and obtain,

$$S^{\text{bulk}} = -\frac{2\gamma}{\ell} M^{d-1} \beta V_{d-1} B(\phi_c)^d e^{-\gamma \int_0^{\phi_c} (X - \frac{Y}{X}) d\tilde{\phi}} \int_{\phi_c}^{\phi_h} d\phi \frac{1 - X^2 + Y}{X} e^{\gamma \int_{\phi_c}^{\phi} d\tilde{\phi} \frac{1 - X^2 + Y}{X}}. \quad (\text{C.37})$$

Integrand is a total derivative, thus

$$S^{\text{bulk}} = -\frac{2}{\ell} M^{d-1} \beta V_{d-1} B(\phi_c)^d e^{-\gamma \int_0^{\phi_c} (X - \frac{Y}{X}) d\tilde{\phi}} e^{\gamma \int_{\phi_c}^{\phi} d\tilde{\phi} \frac{1 - X^2 + Y}{X}} \Big|_{\phi_c}^{\phi_h}. \quad (\text{C.38})$$

This can be simplified further: using (C.4), one realises that the integrand in the exponent is a total derivative of $\log Y(\phi)$. Thus, one has,

$$S^{\text{bulk}} = -\frac{2}{\ell} M^{d-1} \beta V_{d-1} B(\phi_c)^d e^{-\gamma \int_0^{\phi_c} (X - \frac{Y}{X}) d\tilde{\phi}} \left(\frac{Y(\phi_c)}{Y(\phi_h)} - 1 \right). \quad (\text{C.39})$$

But $Y(\phi_h) = \infty$ by regularity condition at the horizon (see section C.29), hence we have the final expression for the Einstein contribution on the BH geometry:

$$S^{\text{bulk}} = \frac{2}{\ell} M^{d-1} \beta V_{d-1} B(\phi_c)^d e^{-\gamma \int_0^{\phi_c} (X - \frac{Y}{X}) d\tilde{\phi}}. \quad (\text{C.40})$$

Gibbons-Hawking contribution. We move on to the Gibbons-Hawking term that is given by the frame-independent expression, the second term in (2.1):

$$S_{\text{GH}} = -2M^{d-1} \int_{\partial M} d^d x \sqrt{h} K \quad (\text{C.41})$$

with

$$K_{\mu\nu} \equiv \nabla_\mu n_\nu = \frac{1}{2} n^\rho \partial_\rho h_{\mu\nu} \quad , \quad K = h^{ab} K_{ab} \quad (\text{C.42})$$

where h_{ab} is the induced metric on the boundary and n_μ is the (outward directed) unit normal to the boundary. In the ϕ -frame (C.31), it is given by

$$n^\mu = -\frac{1}{\sqrt{g_{\phi\phi}}} \left(\frac{\partial}{\partial \phi} \right)^\mu = \frac{\delta^\mu_\phi}{\sqrt{g_{\phi\phi}}}. \quad (\text{C.43})$$

The determinant of the induced metric on the boundary and the extrinsic curvature now are

$$\sqrt{h} = B(\phi_c)^d \sqrt{F(\phi_c)}, \quad (\text{C.44})$$

and

$$K = \frac{\gamma D(\phi_c) \sqrt{F(\phi_c)}}{X(\phi_c) B_0(\phi_c)} \left(1 + \frac{Y(\phi_c)}{2} \right). \quad (\text{C.45})$$

Therefore one finds,

$$S_{\text{GH}} = -\frac{2d}{\ell} M^{d-1} \beta V_{d-1} B(\phi_c)^d F(\phi_c) e^{-\gamma \int_0^{\phi_c} X} \left(1 + \frac{Y(\phi_c)}{2} \right). \quad (\text{C.46})$$

Counter-term. The counter-term action is expressed in terms of the superpotential in equation (B.12). Using equation (C.17) we obtain

$$S_{\text{ct}} = \frac{2(d-1)}{\ell} M^{d-1} \beta V_{d-1} B(\phi_c)^d e^{-\gamma \int_0^{\phi_c} X_{\text{ct}} - \frac{Y}{2X}}, \quad (\text{C.47})$$

where X_{ct} is the analytic solution to (C.3) obtained in terms of odd powers of ϕ . We will only need the first term $X_{\text{ct}} = -X_- \phi + \dots$ in what follows.

The total free energy. The total free energy now can be obtained as the sum of (C.40), (C.46) and (C.47). Using the relation between the free energy and the on-shell gravity action $F = S/\beta$, equations (C.23), (C.22) and (C.28) we find in the limit $\phi_c \rightarrow 0$,

$$F = (M_p \ell)^{d-1} V_{d-1} |j|^{\frac{d}{\Delta_-}} (2(d-1) \operatorname{sgn}(j) X_- \delta X_0(\phi_h) - Y_0(\phi_h)) . \quad (\text{C.48})$$

The physical meaning of the constants δX_0 and Y_0 are as follows. Solving (C.6) near the boundary and matching onto (2.11a) and using (2.12) one finds that the VeV is given in terms of δX_0 as,

$$\langle \mathcal{O} \rangle = \frac{d}{\gamma} |j|^{\frac{d}{\Delta_-} - 1} \delta X_0(\phi_h) . \quad (\text{C.49})$$

On the other hand solving Y from (C.1) near the boundary using the near boundary expansions, and the definitions of T and s from (2.5) we find Y_0 in terms of enthalpy as

$$Y_0(\phi_h) = \frac{T s}{d} (M_p \ell)^{1-d} |j|^{\frac{d}{\Delta_-}} . \quad (\text{C.50})$$

Then the free energy (C.48) can directly be expressed in terms of the enthalpy and the VeV as in (2.23).

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