



# FORWARD: A Toolset for Multiwavelength Coronal Magnetometry

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Determining the 3D coronal magnetic field is a critical, but extremely difficult problem to solve. Since different types of multiwavelength coronal data probe different aspects of the coronal magnetic field, ideally these data should be used together to validate and constrain specifications of that field. Such a task requires the ability to create observable quantities at a range of wavelengths from a distribution of magnetic field and associated plasma—i.e., to perform forward calculations. In this paper we describe the capabilities of the FORWARD SolarSoft IDL package, a uniquely comprehensive toolset for coronal magnetometry. FORWARD is a community resource that may be used both to synthesize a broad range of coronal observables, and to access and compare synthetic observables to existing data. It enables forward fitting of specific observations, and helps to build intuition into how the physical properties of coronal magnetic structures translate to observable properties. FORWARD can also be used to generate synthetic test beds from MHD simulations in order to facilitate the development of coronal magnetometric inversion methods, and to prepare for the analysis of future large solar telescope data.

**Keywords:** sun: corona, sun: magnetic fields, sun: x-rays, sun: radio, sun: infrared, sun: EUV

## 1. INTRODUCTION

In essence, the goal of coronal magnetometry is to solve an inverse problem. Given magnetically-sensitive coronal observations (including, but not limited to polarimetry), the challenge is to determine the magnetic field distribution that generates them. Solving such an inverse problem requires three things: a means of specifying the physical state (e.g., the distribution of density, temperature, velocity, and magnetic field), a well-defined forward calculation (i.e., the physical process relating the physical state and the observations), and the observations themselves.

FORWARD is a set of more than 200 IDL procedures and functions that form a SolarSoft (Freeland and Handy, 1998) package for synthesizing observables and comparing them to coronal data from EUV/Xray imagers, UV/EUV spectrometers, visible/IR/UV polarimeters, white-light coronagraphs, and radio telescopes. It may be called from the command line (i.e., `for_drive`), or via a widget interface (i.e., `for_widget`; Forland et al., 2014). The standard output product is a 2D plane-of-sky map, 2D latitude-longitude (Carrington) map, or user-specified spatial sampling

(**Figure 1**). Image field of view and resolution is user-controlled, as is “viewer” position and line-of-sight (LOS) integration spacing and limits. Details on how to run and install FORWARD are available at <http://www.hao.ucar.edu/FORWARD/>.

This paper describes how FORWARD addresses all three of the requirements for coronal magnetometric inversion and gives examples of how it may be used. Section 2 demonstrates how the physical state may be defined through analytic or numerical models, either user-inputted or generated by FORWARD through included codes or via its interface with online coronal simulations. Section 3 describes the multiwavelength forward calculations that predict observational manifestations of physical processes such as *Thomson scattering*, *collisional excitation*, *continuum absorption*, *resonance scattering*, *Zeeman and Hanle effects*, *Doppler shift*, *thermal bremsstrahlung*, *gyroresonance*, and *Faraday rotation*, and discusses the magnetic diagnostic potential of each. Section 4 describes how FORWARD enables the access and manipulation of observations and converts them to a format directly comparable to the predictions of forward calculations. Section 5 shows how FORWARD may be applied to validate models, build intuition regarding coronal magnetic signatures, tune models to match data, and generally guide the development of multiwavelength magnetometric inversion techniques. Finally, in Section 6 we present our conclusions.

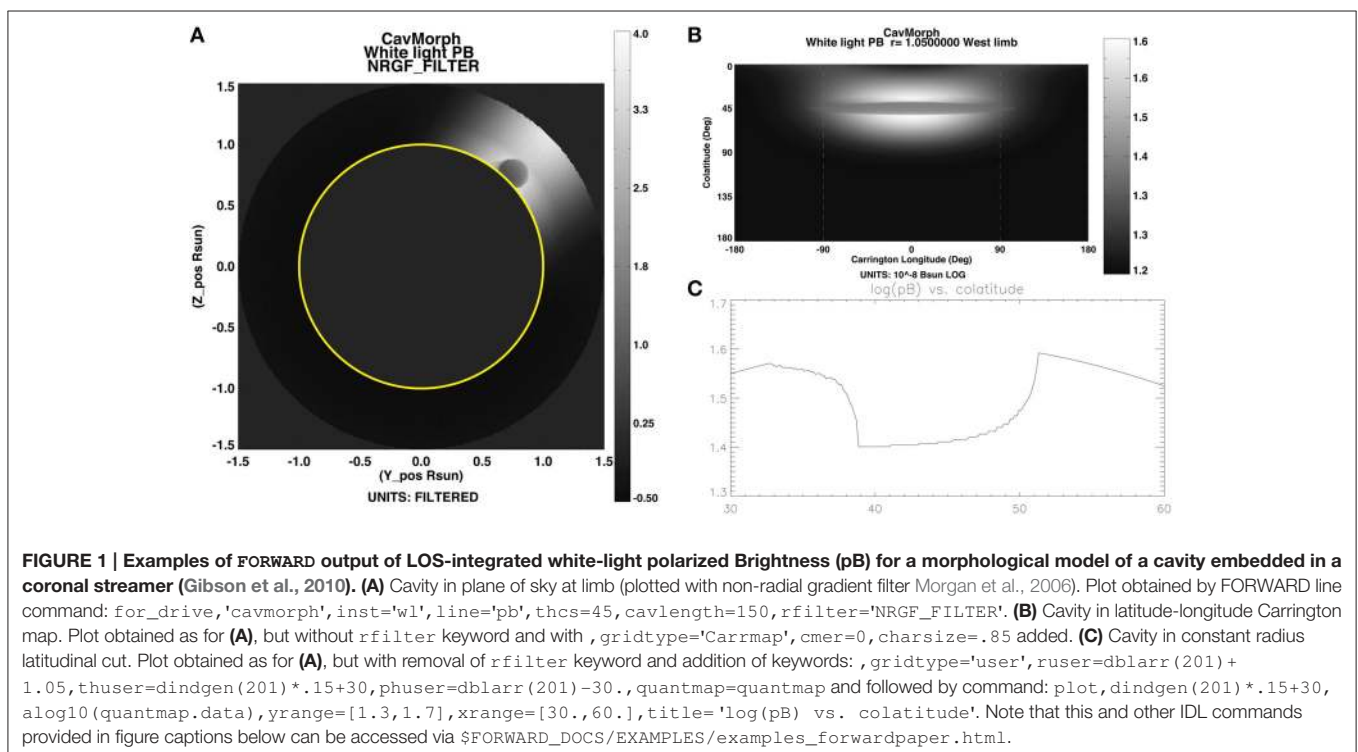
## 2. THE PHYSICAL STATE

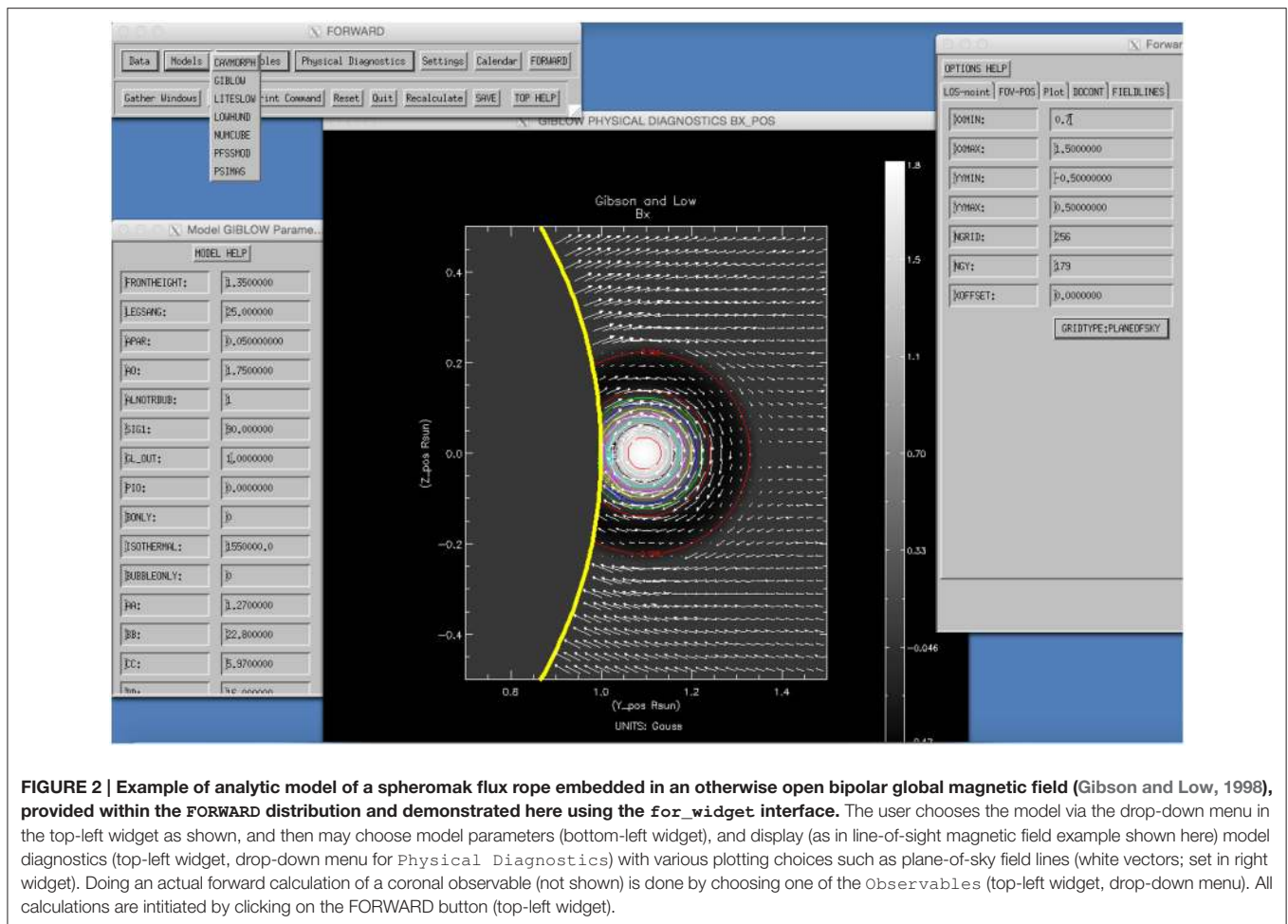
When discussing solar-coronal forward analysis, it is important to differentiate between the model of the *physical state* of the

corona, which addresses the distribution of magnetic fields and plasma throughout 3D space, and the model of how these fields and plasma operate in the presence of a *physical process*, which enables the synthesis of an observed quantity. We will treat the latter in Section 3 as the heart of the forward calculation.

Models of the physical state essentially create synthetic Suns—generally through solutions of the MHD equations. FORWARD includes several analytic models in its distribution (i.e., Low and Hundhausen, 1995; Lites and Low, 1997; Gibson et al., 2010; **Figure 1**; Gibson and Low, 1998; **Figure 2**). It is straightforward to expand it to incorporate other analytic models. Alternatively, a user may input a numerical data cube describing the 3D distribution of plasma and fields. If the data cube is not global, options are provided regarding what to do outside the cube (e.g., zero, constant, or dipolar field, and hydrostatic atmospheres—either isothermal-exponential or power-law). If the data cube only provides a magnetic field, hydrostatic atmospheres can be applied throughout space. (See [http://www.hao.ucar.edu/FORWARD/FOR\\_SSW/idl/MODELS/NUMCUBE/make\\_my\\_cube.pro](http://www.hao.ucar.edu/FORWARD/FOR_SSW/idl/MODELS/NUMCUBE/make_my_cube.pro) for instructions on how to convert a numerical data cube to FORWARD format.) In addition to the forward-calculated observables discussed in Section 3, FORWARD allows easy display of the parameters of the physical state, e.g., density, temperature, magnetic field, velocity (see e.g., **Figure 2**).

Given a calendar date, FORWARD can also automatically interface with the SolarSoft Potential Field Source Surface (PFSS) package (<http://www.lmsal.com/~derosa/pfsspack/>) and





**FIGURE 2 | Example of analytic model of a spheromak flux rope embedded in an otherwise open bipolar global magnetic field (Gibson and Low, 1998), provided within the FORWARD distribution and demonstrated here using the `for_widget` interface.** The user chooses the model via the drop-down menu in the top-left widget as shown, and then may choose model parameters (bottom-left widget), and display (as in line-of-sight magnetic field example shown here) model diagnostics (top-left widget, drop-down menu for `Physical Diagnostics`) with various plotting choices such as plane-of-sky field lines (white vectors; set in right widget). Doing an actual forward calculation of a coronal observable (not shown) is done by choosing one of the `Observables` (top-left widget, drop-down menu). All calculations are initiated by clicking on the `FORWARD` button (top-left widget).

the web-served Magnetohydrodynamic Algorithm outside a Sphere (MAS)-corona MHD simulation data cubes ([http://www.predsci.com/hmi/data\\_access.php](http://www.predsci.com/hmi/data_access.php); Figure 3; Lionello et al., 2009). This enables global descriptions of the 3D coronal magnetic field, and for the MAS model also the plasma in MHD force balance, specific to a given time/day and viewer position.

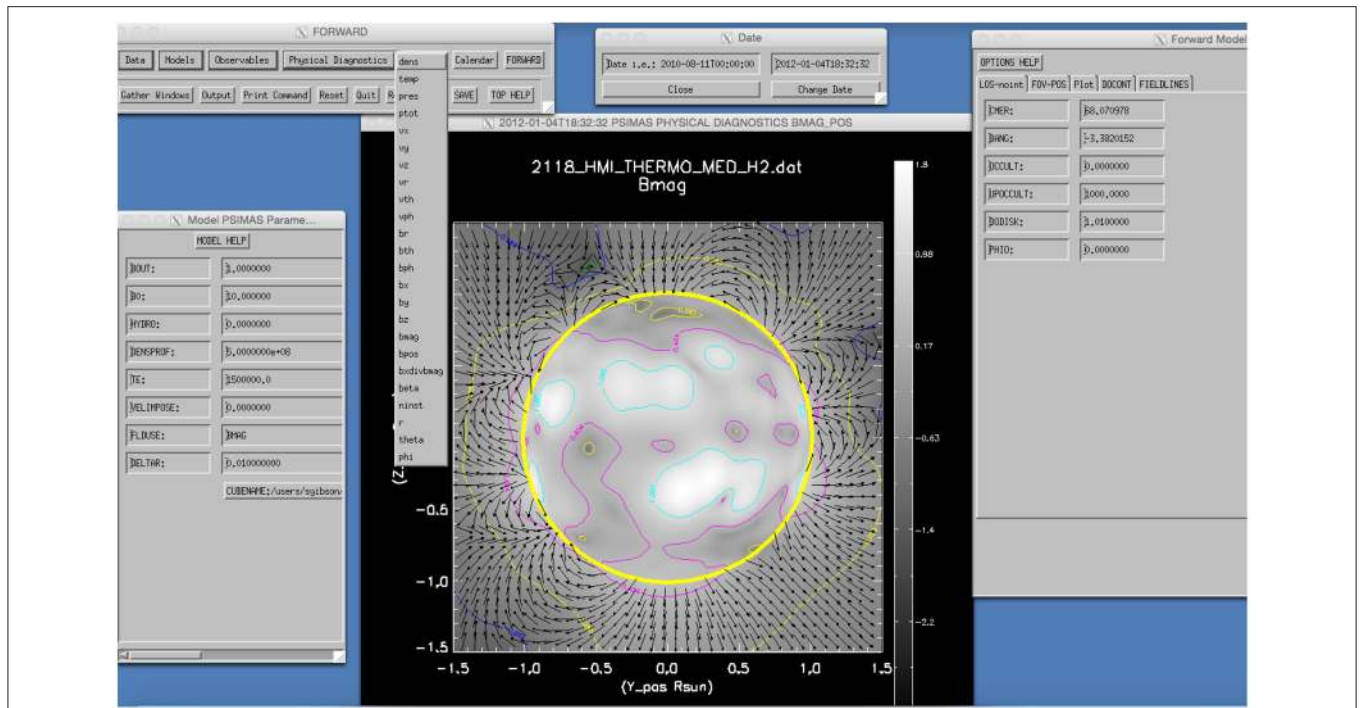
FORWARD also allows the user to specify the physical state of two populations of plasma that may need to be treated independently in the forward calculation. For example, a user may specify a population of plasma at a coronal temperature, and another population of cooler, chromospheric plasma subject to continuum absorption in EUV images (see Section 3.3). This provides capability, for example, for depicting models where cool solar prominences exist in the context of surrounding coronal temperature material, such as those produced by Luna et al. (2012) or Xia et al. (2014) (see also the cavity-prominence test-bed simulation shown in Section 5). Another application would be to allow models where two different coronal populations lie along the line of sight, each with different abundance properties. It is also possible to set a filling factor for one or both populations. This capability allows exploitation of the diagnostic potential of comparing emissions which may have different dependencies on density (as we discuss in Section 3).

### 3. THE FORWARD CALCULATION: PHYSICAL PROCESSES

Given a modeled physical state, i.e., a specification of the distribution of density, temperature, magnetic field and velocity in the corona, FORWARD is able to produce many different synthetic observables. These observables arise from various physical processes manifesting at different wavelengths of light in the corona. They depend upon the viewer's line of sight, along which (for example) optically-thin emission must be integrated. FORWARD establishes these lines of sight either through keyword definition of an observer's heliographic latitude and longitude, or through keyword setting of a calendar date from which the position of the Earth (or STEREO spacecraft) can be determined. In this section, we will discuss a range of physical processes relevant to the corona, describe how they translate to observables that FORWARD synthesizes, and consider their potential for coronal magnetometry. Table 1 provides a summary.

#### 3.1. Thomson Scattering

Thomson scattering is the main physical process responsible for illuminating the continuum, or "K" corona. Photospheric



**FIGURE 3 | Example of MAS model data cube loaded into FORWARD by date, and displaying magnetic field strength (contours) and field lines (black arrows) in the plane of the sky.** As with the analytic model of Figure 2, display options and model parameters can be set through the widgets. Descriptions of the model parameters and other options are found via the TOP HELP, MODEL HELP, and OPTIONS HELP buttons at the top of the three main widget windows. See Section 3 for further discussion and Figures 4, 6 for examples of forward calculation using the MAS and PFSS models.

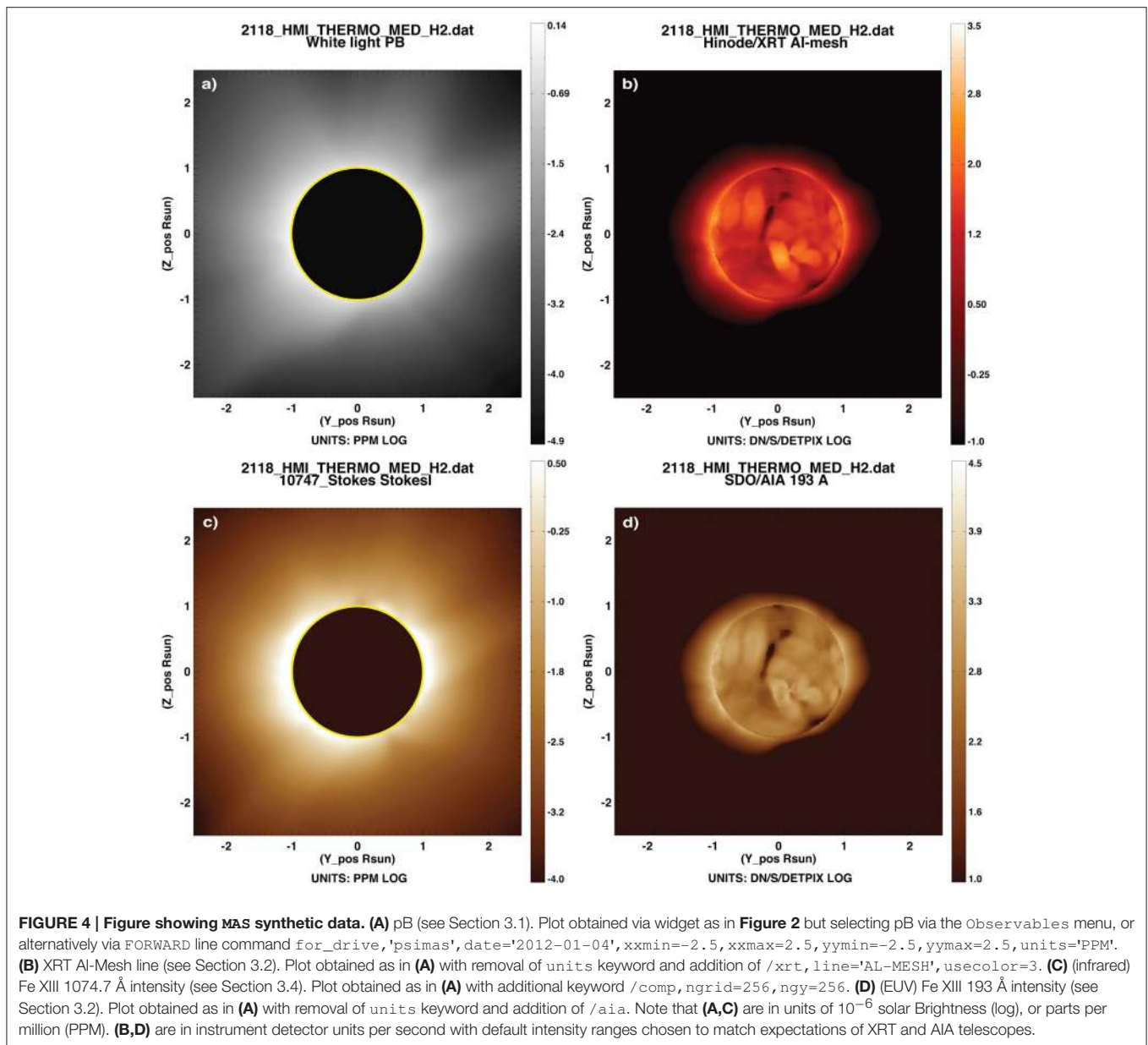
**TABLE 1 | Physical processes as defined in Section 3, highlighting dependency on attributes of the physical state, which observations are sensitive to them, and diagnostic sensitivity to the 3D coronal magnetic field.**

Process	Physical-state dependency	Observation	Magnetic quantity probed
Thomson scattering	Electron density	White-light $pB$ , $TB$	Plasma structured by field (e.g., closed vs. open field boundaries, flux surfaces)
Collisional excitation	Electron density, temperature	IR/Visible/EUV/SXR emission	Plasma structured by field (incl. loops, closed/open boundaries, flux surfaces)
Continuum absorption	Chromospheric population density, electron density, temperature	EUV absorption features	Can indicate magnetic geometry suitable for prominence formation
Resonance scattering; polarization	Electron density, temperature, vector magnetic field	Visible/IR spectra	$B_{los}$ from Stokes V; Magnetic field direction from Stokes Q, U
Doppler shift	Electron density, temperature, velocity	Visible/IR spectra	$B_{pos}$ and field line direction from waves; flux surfaces from bulk flows
Thermal bremsstrahlung	Electron density, temperature, vector magnetic field	Radio emission (intensity and circular polarization) as a function of frequency	$B_{los}$ from Stokes V
Gyroresonance	Electron density, temperature, vector magnetic field	Radio emission (intensity and circular polarization) as a function of frequency	Surfaces of constant magnetic field strength at each frequency
Faraday rotation	Electron density, temperature, vector magnetic field	Rotation of plane of polarization	$B_{los}$ from rotation measure

light scatters off of free coronal electrons and results in both unpolarized and linearly polarized emission. Both the total brightness ( $TB$ ) and polarized brightness ( $pB$ ) of white light are proportional to  $n_e$  and to a scattering function that depends upon radial distance from the photosphere (Billings, 1966). They are also integrated along the line of sight in the optically-thin corona.

Given a distribution of electron density, FORWARD can synthesize images of  $TB$ ,  $pB$ , and degree of polarization  $p$  (Figures 4A, 5A), comparable to observations from white light coronagraphs such as SOHO/LASCO, STEREO/SECCHI, and MLSO/KCOR. If keyword  $for$  is set, FORWARD will call upon SolarSoft function  $forpol\_KL.pro$  in order to add a model





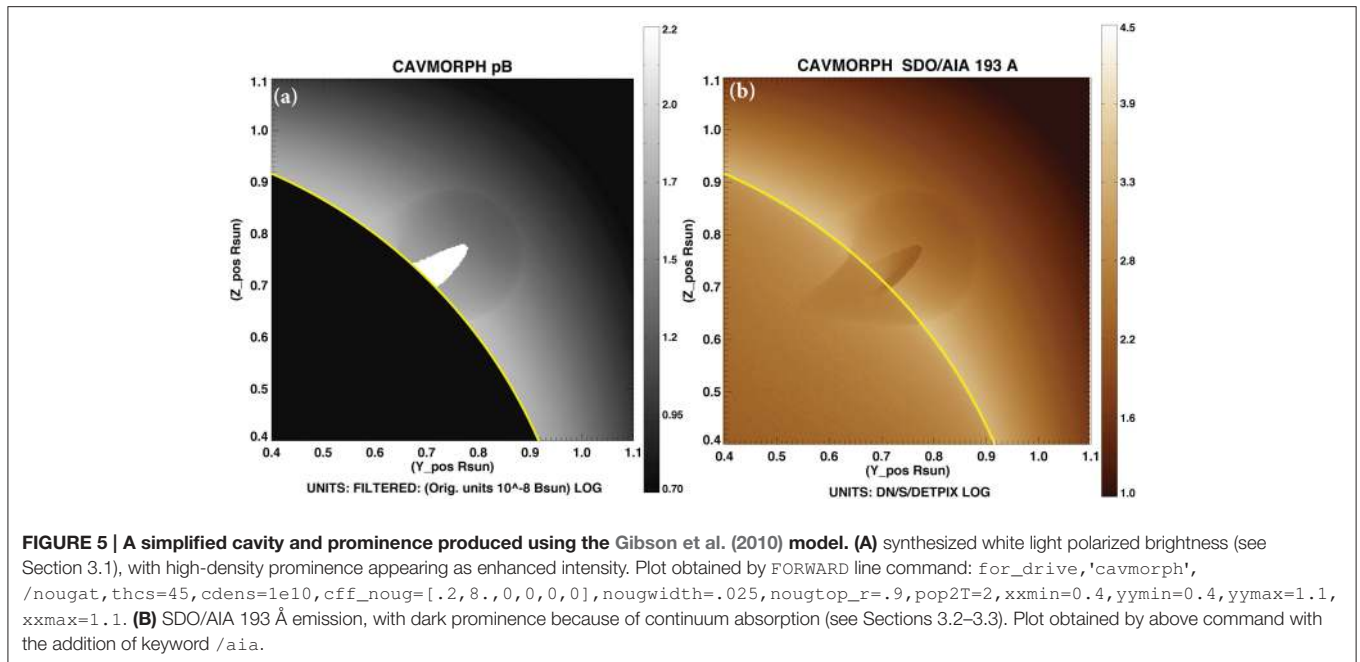
distribution of F-coronal brightness (Koutchmy and Lamy, 1985). This arises from light diffracting through interplanetary particles in the plane of the ecliptic, and is also known as the zodiacal light. It is essentially unpolarized in the first few solar radii (Mann, 1992).

Thomson scattering has no direct dependency on magnetic field, but there is sensitivity to magnetic topology through its dependence on density. For example, bright (dense) coronal streamers generally correspond to closed magnetic fields, and dark (sparse) coronal holes generally correspond to open magnetic fields. For this reason, white light coronagraph data have been used to qualitatively validate features of coronal magnetic models (Newkirk and Altschuler, 1970), and more quantitatively, to define the average nonradial expansion of

magnetic fields in coronal holes (Kopp and Holzer, 1976; Munro and Jackson, 1977). Magnetic flux surfaces also may be delineated by white light structures, such as three-part CME features (Low and Hundhausen, 1995), and prominence cavities (e.g., **Figure 5A**; see also Gibson and Fan, 2006).

### 3.2. Collisional Excitation

Solar coronal radiation in the extreme ultraviolet (EUV) and soft X-ray (SXR) is produced by collisionally-excited atoms in thermal and ionization equilibrium. The intensity of this emission is proportional to  $n_e^2$  and the temperature response of the line(s). For spectrographs, FORWARD calculates the integrated intensities of lines in physical units to be compared with processed spectral data. For waveband imagers, FORWARD



incorporates the wavelength-response function of the instrument into its calculated intensities. The emission is integrated along the line of sight in the optically-thin corona.

FORWARD synthesizes images comparable to those produced by numerous EUV and Soft X-ray imagers and also many spectral line intensities in these wave bands. Currently simulated imagers include SOHO/EIT, STEREO/EUVI, Hinode/XRT (e.g., **Figure 4B**), ProbA-2/SWAP, and SDO/AIA (e.g., **Figures 4D, 5B**). Adding new imagers is straightforward if the wavelength response function is available. Count rates for imagers are calculated by convolving the wavelength response function of the imager with pre-calculated spectra at various temperatures and densities, produced using the Chianti atomic data base and related software (Dere et al., 1997; Del Zanna et al., 2015). For the imagers, users may select from pre-calculated abundance options; current selections include coronal abundances determined by Feldman et al. (1992) or Schmelz et al. (2012) or photospheric abundances of Caffau et al. (2011). The code uses the Chianti ionization equilibrium calculations (Dere et al., 2009).

Spectral line intensities can be calculated for any line between 1 and 1410 Å, again under the assumption that the coronal plasma is collisionally excited and in thermal and ionization equilibrium. Based on a user-specified instrumental line-width, the code includes any blended lines included in the Chianti spectral line calculations. For spectral lines, users may specify any abundance or ionization table in the Chianti database format. Default line widths are provided for particular instruments like Hinode/EIS, IRIS, and SoHO/CDS, but the code is not limited to lines observed by these instrument. The wavelength range covers the IRIS far ultraviolet (FUV) range, but not the near ultraviolet (NUV) range, which does not include lines that can be modeled by FORWARD.

As with Thomson scattering, radiation from collisional excitation does not have a direct dependence on magnetic fields. However, the suppression of conductivity across magnetic field lines means that magnetic field lines are essentially traced out in coronal emission. EUV and SXR structures thus often provide a diagnostic of local magnetic field geometry—for example when coronal loops are lit up in active regions (see e.g., Savcheva et al., 2013; Malanushenko et al., 2014; also Savcheva and Malanushenko, in preparation). Magnetic boundaries or flux surfaces are also often delineated due to sharp density/temperature gradients, as in the case of open vs. closed fields (e.g., **Figures 4B,D**), and prominence cavities (e.g., **Figure 5B**).

### 3.3. Continuum Absorption

Relatively cool, chromospheric temperature material suspended in the corona (e.g., a prominence) results in Lyman continuum absorption by neutral hydrogen and by neutral and once-ionized helium (see Kucera, 2015 for further details).

The observed intensity is

$$I = I_f + I_b e^{-\tau} \quad (1)$$

where  $I_f$  is the foreground radiation,  $I_b$  is the background radiation, and  $\tau$  is the continuum absorption summed over the three absorbing species:

$$\tau = \sum_i \sigma_i \int n_i dh \quad (2)$$

where  $h$  is the distance along the line of sight,  $n_i$  is the number density of each species, and  $\sigma_i$  is the absorbing cross section as a function of wavelength (values calculated with the formulation of Keady and Kilcrease, 2000).

Through the definition of a second population of low-temperature plasma with a specified density distinct from the primary coronal population, FORWARD calculates the effect of continuum absorption on the total intensity in EUV. **Figure 5B** shows a model of a simple prominence inside a cavity in the 193 Å band of SDO/AIA. The cavity is darker than its surroundings because it has a lower density, but the central prominence is darker because of continuum absorption of background emission.

Again, continuum absorption has no direct dependence on magnetic fields. However, magnetic field geometry (e.g., dipped or flat field lines) is expected to play an important role in establishing where prominences form (see Karpen, 2014 and references therein).

### 3.4. Resonance Scattering and Polarization

Emission from the coronal forbidden lines arises both from collisional excitation as described above in Section 3.2, and also from resonance scattering. In resonance scattering, anisotropic radiation from the underlying photosphere excites coronal ions and leads to reemitted light with a characteristic polarization signature (Casini and Judge, 1999). This emission depends linearly upon ion density (and thus electron density), as opposed to quadratically as in collisional excitation. **Figure 4C** vs. **Figure 4D** illustrates the difference between the collisionally-excited/resonantly-scattered infrared Fe XIII line and the predominantly collisionally-excited EUV Fe XIII line. Intensity of the latter drops off quickly, while the former shows similarity in the outer field of view to the Thomson-scattered white light (**Figure 4A**), which is also linearly-dependent upon electron density at these heights (see Habbal et al., 2011 for further discussion).

FORWARD employs the Coronal Line Emission (CLE) Fortran-77 polarimetry code developed by Judge and Casini (2001) to synthesize Stokes ( $I$ ,  $Q$ ,  $U$ ,  $V$ ) line profiles for the visible and infrared forbidden lines including Fe XIII 1074.7 and 1079.8 nm (currently observed by MLSO/CoMP as discussed in Section 4), Fe XIV 530.3 nm, Si IX 393.4 nm, and Si X 1430.5 nm. Stokes  $I$  indicates the total intensity of the line,  $Q$  and  $U$  together constitute its linearly polarized intensity, and  $V$  is the circularly polarized intensity. The CLE code models the lines under the combined influence of resonance scattering and particle collisions in the presence of coronal magnetic fields.

Because of its sensitivity to magnetic fields, the Stokes  $I$ ,  $Q$ ,  $U$ ,  $V$  polarization vector can be used as a direct diagnostic of coronal magnetism (subject to intensity-weighted line-of-sight integration). For example, the *Zeeman effect* generates circularly polarized light (Stokes  $V$ ) proportional to line-of-sight-oriented magnetic field  $B_{los}$ . Since the coronal visible/infrared forbidden lines treated by CLE have a Larmor frequency  $\nu_L \approx \mu_B B/h$  that is much larger than the inverse lifetime of the atomic transitions being modeled, they lie in the strong field (or saturation) limit of the *Hanle effect* (see Raouafi et al., submitted; Dima et al., submitted) for discussion of magnetometry in the UV “unsaturated” Hanle regime). In the saturated regime, linear polarization provides a probe of the direction of the magnetic

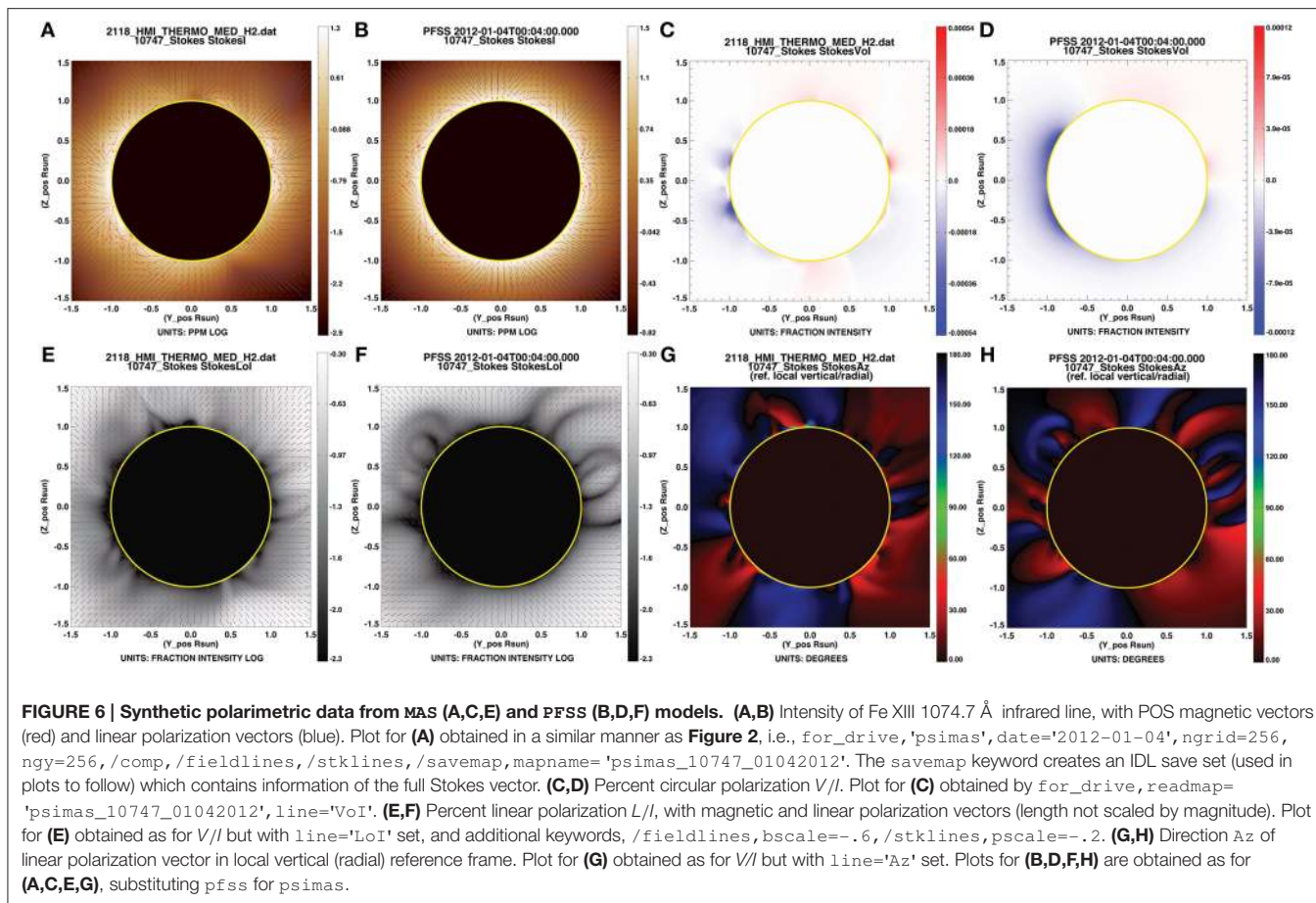
field in the plane-of-sky (POS), but not its strength. In particular, the direction of the linear polarization vector [or azimuth,  $Az = -0.5 * \text{atan}(U/Q)$ ] is parallel to the POS component of the magnetic field, as long as the local magnetic vector field has an angle relative to the solar radial direction ( $\vartheta_B$ ) less than the critical “*van Vleck*” angle, at which point the azimuth becomes perpendicular to the POS field. This occurs because the atomic alignment upon which the linear polarization depends goes through zero (and changes sign) when  $3\cos^2(\vartheta_B) = 1$ , i.e., when  $\vartheta_B = 54.74^\circ$  (van Vleck, 1925). The location of van Vleck nulls in linear polarization  $L = \sqrt{Q^2 + U^2}$  thus also acts as a diagnostic of magnetic field direction (see Section 5.1 for further discussion).

**Figures 6A,C,E,G** shows  $I$ ,  $V/I$ ,  $Az$  and  $L/I$  calculated from the MAS coronal model. From this it is clear that, despite the line-of-sight superposition of optically-thin coronal plasma, Stokes polarimetry can provide a quantitative measure of coronal magnetic field strength and direction. The Stokes  $V/I$  (**Figure 6C**) represents a line-of-sight intensity-weighted average of  $B_{los}$ . The dark linear-polarization features shown in **Figure 6E**) are generally signatures of magnetic fields oriented at the van Vleck angle (although note that the presence of strong  $B_{los}$  field can also result in linear polarization nulls; see Section 5.1 for further discussion). Even with LOS integration, the linear polarization vectors (blue) are largely aligned with the POS magnetic field vectors (red) (see **Figures 6A,E**), except when at the van Vleck angle they flip  $90^\circ$ . **Figure 6G**) also illustrates this sensitivity to POS magnetic field direction, showing magnitude of departure from radial-orientation in  $Az$  (red = counterclockwise, blue = clockwise). The coronal hole in the south/southwest is evident as a broad blue/red interface in  $Az$ , indicative of diverging magnetic fields, while closed field structures exhibit a red-black-blue interface (e.g., south/southeast) which indicates converging fields.

Strongly nonradial azimuths (represented as green in the local-vertical reference frame of **Figure 6G**) are rare. They can occur if the local magnetic vector  $> 54.74^\circ$  as measured from the solar radial direction but the POS projection is close to radial, as in the case of magnetic fields that are oriented largely along the LOS. In order for such a nearly-perpendicular azimuth to survive LOS integration, either the plasma must be localized to a magnetic structure oriented in this manner, or a larger-scale magnetic structure must possess a symmetry along the LOS. Such symmetries are fairly common in large-scale POS-oriented fields extended along the LOS, e.g., arcade fields or coronal holes, and because such structures are POS-oriented, they possess a strong linear-polarization signal and so the azimuth survives LOS integration (see further discussion in Section 5.1). Even if LOS-oriented fields are localized or exist with orientation extended along the LOS, however, because they do not possess a strong linear polarization signal they are likely to be obscured by any POS-oriented fields lying along their integration path.

**Figures 6B,D,F,H** shows the polarization for a potential field model extrapolation for the same day and using similar (although not identical) photospheric magnetic boundary data as the MAS model of **Figures 6A,C,E,G**. The differences between MAS and





PFSS predictions for circular and linear polarization result from differences at the lower boundary, from the non-potentiality of the MAS model magnetic field, and also to some degree from differences in intensity-weighting along the line of sight (the PFSS solution requires a density/temperature distribution that is spherically-symmetric). The significance of intensity weighting is also evident in **Figure 7**, where the LOS-integrated Stokes  $V$  differs depending on the wavelength used to observe it (visible, IR, radio). Since the same (MAS) model is used for all four forward calculations, variation must be due to the different sensitivities to temperature and density for the four wavelength regimes, which in turn means that different distributions of plasma are contributing to the integrals along the line of sight. In Section 5 we will discuss the importance of making full use of such multiwavelength magnetic dependencies in choosing between models.

### 3.5. Doppler Shift

If light-emitting plasma is moving, spectral lines are subject to a Doppler shift proportional to the line-of-sight component of the plasma velocity ( $v_{los}$ ). For optically-thin plasma, this  $v_{los}$  is further weighted by the distribution of intensity along the line of sight. From the line profiles in the visible and IR generated by CLE (see Section 3.4), FORWARD determines Doppler shift and integrates

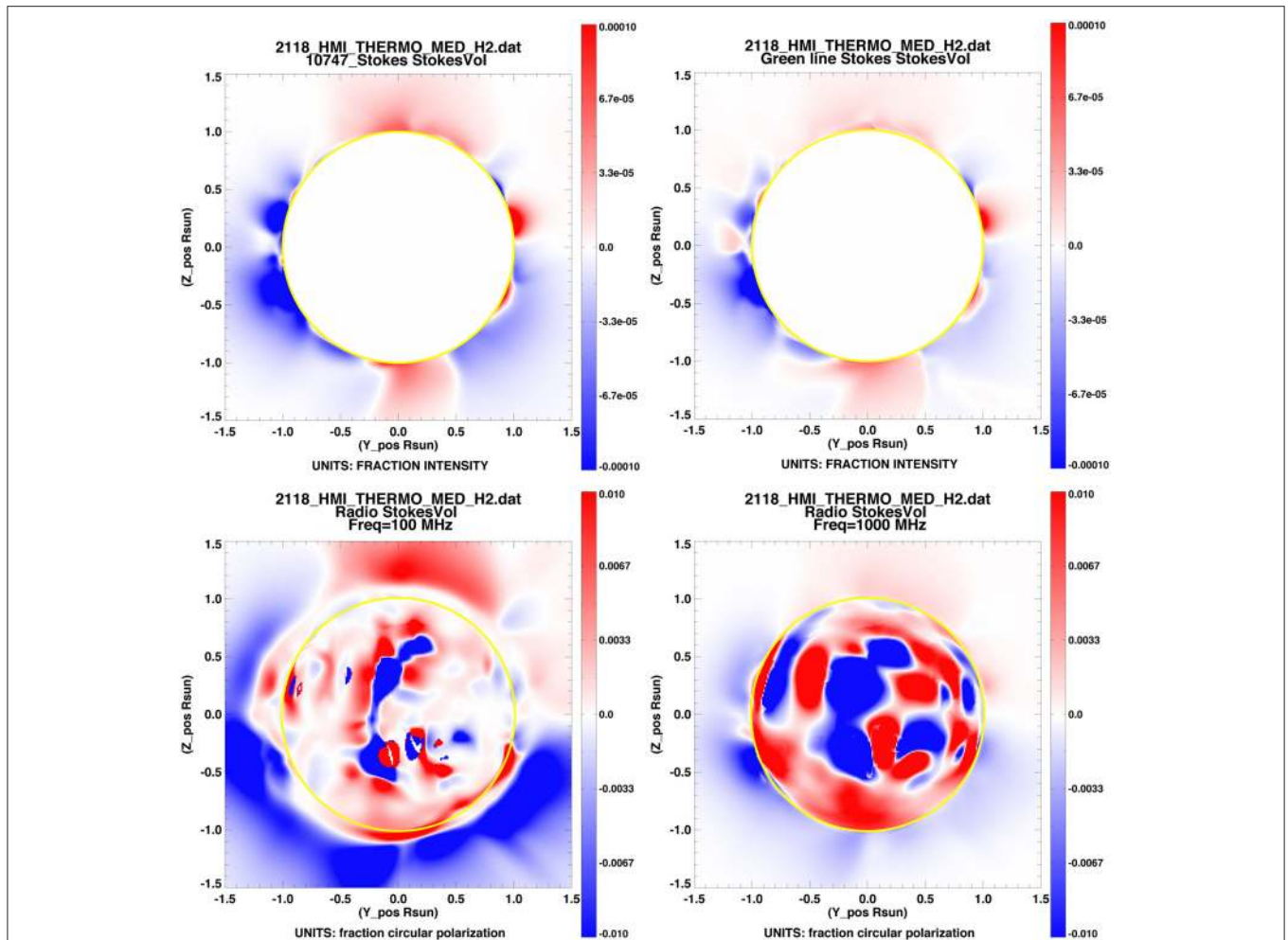
along the line of sight to get a synthetic observable comparable to observations.

Doppler velocity observations in the IR by the MLSO/CoMP telescope have proved to be a good resource for measuring ubiquitous waves in the corona (Tomczyk et al., 2007). The phase speeds of these waves are expected to be proportional to the plane-of-sky component of the magnetic field strength, and the direction of propagation of the waves will be aligned with the magnetic field direction. In general, the flux-freezing condition forces plasma flows to follow the direction of the magnetic field, so bulk velocity flows also can act as a probe of magnetic structure (**Figure 10C**; Bąk-Stęślicka et al., 2013, also Bąk-Stęślicka et al., submitted).

### 3.6. Radio Emission: Thermal Bremsstrahlung and Gyroresonance

The two thermal emission mechanisms that dominate non-flaring solar radio emission are bremsstrahlung (also known as “free-free emission”) and gyroresonance emission. Bremsstrahlung is produced by all plasma in the solar atmosphere and is strongest in dense regions, while gyroresonance emission requires strong magnetic fields in the corona and is usually confined to locations above sunspots. The Jansky Very Large Array, the Expanded Owens Valley Solar Array, the





**FIGURE 7 |** Because of different dependencies on plasma along the line of sight, the representation of integrated circular polarization (dependent upon line-of-sight magnetic field strength) appears differently at different wavelengths. (Top left) *MAS* model *V//* for Fe XIII 1074.7. Plot obtained as in Figure 6, with additional keywords `imin=-0.00001, imax=0.00001`. (Top right) Same for Fe XIV green line. Plot obtained by following the process outlined in Figure 6, but substituting `/greencomp` for `/comp`. (Bottom left) Same for radio bremsstrahlung, at a frequency of 100 MHz. Plot obtained as in Figure 6, but with `/radio` instead of `/comp` and `frequency_MHz=100, imin=-0.001, imax=0.001`. (Bottom right) Same, but substituting `frequency_MHz=1000`. Unlike the visible and IR lines, radio frequencies can be observed above the solar disk as well as at the limb.

Nobeyama Radioheliograph and the Mingantu Ultrawide Spectral Radioheliograph are examples of radio telescopes capable of high-resolution, high-dynamic-range imaging, including circular polarization imaging, in the frequency range (1–20 GHz) where these two mechanisms are important diagnostics of the magnetic field in the solar atmosphere.

Optical depths are generally significant in the solar atmosphere at radio wavelengths and in order to calculate the radio emission arising from either of these physical processes one must carry out a radiative transfer calculation (as for continuum absorption in Section 3.3). It is convenient to do the calculation in terms of brightness temperature,  $T_B$ , because radio emission takes place in the Rayleigh–Jeans limit where the effective radiative temperature of an optically thick source is the physical temperature of that source. Brightness temperature may be converted to flux density  $S$  via the relation

$$S = k_B \frac{f^2}{c^2} \int T_B d\Omega \quad (3)$$

where  $k_B$  is Boltzmann’s constant,  $c$  is the speed of light and the integral is over the solid angle  $\Omega$  of interest. Note that brightness temperature is a local quantity whereas flux density is integrated over a source area.

The radiative transfer calculation for radio emission in FORWARD follows standard methods: the brightness temperature transfer is governed by the differential equation (e.g., Dulk, 1985):

$$\frac{dT_B}{ds} = \kappa (T_e - T_B) \quad (4)$$

where  $\kappa$  is the opacity per unit distance  $s$  along the line of sight and  $T_e$  is the local electron temperature. We solve radiative

transfer by determining  $\kappa$  and  $T_e$  in each pixel along the line of sight and integrate Equation 4 across each pixel as follows:

$$T_B' = T_B e^{-d\tau} + T_e (1 - e^{-d\tau}) \quad (5)$$

where  $T_B$  is the incident brightness temperature, and  $T_B'$  is the emergent brightness temperature—integrated across the line-of-sight pixel and serving as the incident brightness temperature to the next pixel.  $d\tau = \kappa ds$  is the opacity change across the pixel.

Radio emission from the solar atmosphere is strongly influenced by the magnetic field in the emitting regions and provides valuable diagnostics of solar magnetic fields that complement other techniques. The magnetic field plays a role in the absorption coefficients  $\kappa$ : electrons interact more strongly with the sense of circular polarization that matches the sense of rotation of an electron as it spirals along magnetic field lines under Larmor motion. The polarization that interacts more strongly with electrons is the *extraordinary* or *x* mode, with the other polarization being labeled the *ordinary* (*o*) mode. Under most conditions in the solar corona, and following propagation to terrestrial observers, the *x* and *o* modes are 100% circularly polarized with opposite sense of polarization. FORWARD solves the radiative transfer equations as described above for each of the circular polarizations separately. The difference between the *x* and *o* modes is then Stokes V (modulo a sign), while the sum is the total intensity, Stokes I (For radio emission from the solar atmosphere, we may ignore any weak linear polarization present due to the fact that the large Faraday rotation in the solar atmosphere wipes out linear polarization over a finite observing bandwidth, see below).

For thermal bremsstrahlung, which is always included in a FORWARD radio emission calculation, opacity results from collisions between electrons and ions. We use the simple expression (Dulk, 1985; Gelfreikh, 2004)

$$\kappa = 0.2 \frac{n_e^2}{T_e^{1.5} (f \pm f_B |\cos \theta|)^2} \quad (6)$$

which is appropriate for coronal temperatures, where  $f_B = 2.8 \times 10^6 B_{\text{gauss}}$  Hz is the electron gyrofrequency and the factor in parentheses deals with polarization (with the assumption that  $f \gg f_B$ ):  $\theta$  is the angle between the magnetic field direction and the line of sight, and the minus sign refers to the *x* mode while the plus sign refers to the *o* mode. Thus, the magnetic field information present in bremsstrahlung emission resides in the circular polarization and represents the line-of-sight component of  $B$ .

The dependencies in Equation 6 mean that bremsstrahlung is strongly favored in dense regions of the atmosphere and weighted toward cooler material (since in Equation 4,  $\kappa T_e \propto T_e^{-0.5}$ , e.g., White, 2000). The  $f^{-2}$  dependence of bremsstrahlung opacity also means that optical depth decreases rapidly as frequency increases, and at low frequencies one is likely to be optically thick such that the lower the frequency, the higher in the atmosphere one sees. This is evident in **Figure 8**, where polarization extends much higher above the photosphere at 100 MHz (lower left panel) than at 1000 MHz (lower right panel). When optically thick,

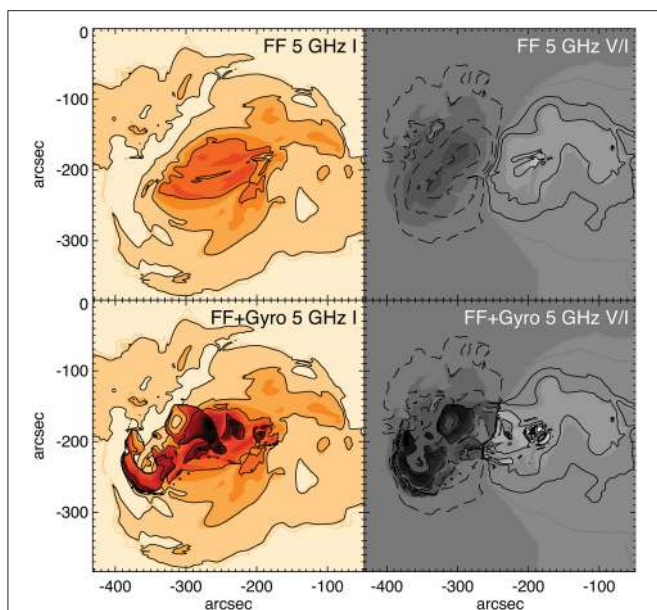
the circular polarization produced by bremsstrahlung emission actually depends on the presence of a temperature gradient. If one has well-calibrated brightness temperature measurements across a continuous frequency range, one can in fact determine both the temperature gradient and the magnetic field from the data (Grebinskij et al., 2000, disproving a comment in White, 2000). In **Figure 7** the lower degrees of polarization over much of the disk at 100 MHz reflect the fact that the temperature gradient is weaker higher in the corona.

The gyroresonance calculation is more complex. Gyroresonance opacity results from the acceleration of electrons in a magnetic field under the Lorentz force, and is only significant in narrow layers where the observing frequency  $f$  is a low integer multiple  $s$  of the electron gyrofrequency  $f_B$  (e.g., White and Kundu, 1997). The optical depth  $\tau$  of a thermal gyroresonance layer (the absorption coefficient integrated through the layer) is

$$\tau_{x,o}(s, f, \theta) \propto \frac{n_e L_B(\theta)}{f} \frac{s^2}{s!} \left( \frac{s^2 \sin^2 \theta}{2\mu} \right)^{s-1} F_{x,o}(\theta) \quad (7)$$

where  $L_B(\theta)$  is the scale length of the magnetic field ( $B/\frac{\partial B}{\partial l}$ ) evaluated along the line of sight and  $\mu = m_e c^2 / k_B T_e$ . For coronal conditions  $\mu \approx 2000$ , and the  $\mu^{-s}$  dependence in Equation 7 produces a dramatic change in opacity as harmonic number  $s$  changes.  $F_{x,o}(\theta)$  is a function of angle which is of order unity for the *x* mode near  $\theta = 90^\circ$ , but decreases sharply at smaller  $\theta$ , and is smaller in the *o* mode than in the *x* mode. FORWARD uses a more exact approximation for  $\tau_{x,o}$  due to Robinson and Melrose (1984) which requires a careful calculation of the cold plasma properties of the electromagnetic modes under the conditions that apply in the gyroresonance layer. FORWARD incorporates gyroresonance emission by testing for harmonic layer crossings along the line of sight in the range  $s = 1$  to 5, and calculating the resulting opacity as shown above: significant gyroresonance opacity at higher harmonics generally requires mildly relativistic electrons which puts emission in the gyrosynchrotron limit in which harmonics are much broader and Equation 7 is no longer valid [Note the simulation package *GX\_Simulator* can handle gyrosynchrotron emission (Nita et al., 2015)].

Gyroresonance emission is most commonly seen in the strong magnetic fields above solar active regions. At frequencies above a few GHz, bremsstrahlung does not produce enough opacity to make the corona optically thick, while the large change in gyroresonance opacity as  $s$  decreases (typically a factor of order 1000) means that a given harmonic layer is usually either very optically thick or very optically thin. When optically thick, gyroresonance produces million-K coronal brightness temperature features in radio images. In practice, we see down to the highest optically thick layer (usually  $s = 3$  in *x* mode and  $s = 2$  in *o* mode), and the brightness temperature variations across the surface (of constant field strength for a given frequency) represent actual temperature variations across that surface. Thus, for this mechanism the magnetic field information contained in the emission morphology is somewhat complex and does not simply reside in the polarization (White and Kundu, 1997).



**FIGURE 8 | Radio emission from a model active region comparing calculations with and without gyroresonance emission.** Thermal bremsstrahlung or free-free emission (“FF”) is included in both cases. Total intensity (Stokes  $I$ ) is plotted in the left panels and degree of polarization ( $V/I$ , scaled from  $-1$  to  $1$ ) is plotted in the grayscale in the right panels, with contours showing Stokes  $V$ . The brightness temperature display range in the left-hand panels is  $0$  to  $2.5 \times 10^6$  K. Contour levels are at  $16$ ,  $32$ ,  $64$ ,  $128$ , and  $256 \times 10^4$  K in the Stokes  $I$  panels, and  $1$ ,  $4$ ,  $16$ , and  $64 \times 10^4$  K in the  $V/I$  images. The pixel size is  $1.5$  arcsec.

Internally, FORWARD carries out radiative transfer for radio emission in the  $x$  and  $o$  modes by summing the requested absorption coefficients in each pixel: bremsstrahlung is the default opacity and gyroresonance opacity may be turned on or off. The brightness temperatures in the two modes are summed and differenced to report Stokes  $I$  and  $V$ , or one can display  $V/I$ , the degree of circular polarization, as in **Figures 7, 8**. Mode coupling between the  $x$  and  $o$  modes, which can result in reversal in the sense of circular polarization at points where the magnetic field direction along the line of sight reverses (e.g., White et al., 1992), is not yet included in the FORWARD calculation but will be in future releases. **Figure 8** shows an example of a FORWARD radio emission calculation using a three-dimensional hydrodynamic active region model (density, vector magnetic field and temperature) with thermal conduction and radiative cooling (Lionello et al., 2013). The upper panels show the model radio emission obtained with just bremsstrahlung opacity included, while the lower panels include gyroresonance opacity. The brightness temperatures are much higher when gyroresonance opacity is included, and the polarization structure becomes more complex.

FORWARD does not currently include plasma emission: this is the dominant emission in low-frequency solar radio bursts (e.g., Kundu, 1965), but it is a coherent emission mechanism and there is no simple way to calculate it (e.g., see Schmidt and Cairns, 2012a,b, for a detailed calculation). In addition, fundamental plasma emission occurs at frequencies where the

refractive index may be significantly different from unity and refraction can play a major role in determining ray paths. FORWARD assumes linear ray paths along lines of sight and does not currently handle refraction at low radio frequencies which will produce curved ray paths in a realistic solar atmosphere.

### 3.7. Faraday Rotation

In general radio emission can be elliptically polarized. Electromagnetic radiation in a magnetized plasma can be decomposed into two natural modes with orthogonal polarizations, and as radiation propagates the two intrinsic polarizations have different refractive indices and slightly different phase speeds. This effect causes the plane of linear polarization to rotate, with the amount of rotation being a function of frequency. In the lower regions of the solar atmosphere the rotation is so large that, as mentioned in the previous subsection, when averaged across a finite observing bandwidth the linear polarization is washed out. However, further out in the solar wind where the magnetic field is lower, Faraday rotation can be measured as a function of frequency, and such measurements are one of the few techniques that can be used as a remote probe of the magnetic field in the solar wind. This technique has been applied to both communication transmissions from satellites (e.g., Bird et al., 1985; Jensen et al., 2005, 2013) as well as polarized background cosmic sources (e.g., Bird et al., 1980; Mancuso and Spangler, 2000; Ord et al., 2007; You et al., 2012; Kooi et al., 2014). In FORWARD it can be used to simulate the contributions of CME and solar wind plasma to an observable diagnostic of the magnetic field.

The expression for Faraday rotation is relatively straightforward: the angle of rotation  $FR$  (in radians) at wavelength  $\lambda$  is

$$FR = RM \lambda^2 \quad (8)$$

where the rotation measure  $RM$  (measured in radians per square meter) is the wavelength-independent measure of Faraday rotation, calculated from the integral of the product of electron density and line-of-sight magnetic field along the ray path:

$$RM = 2.6 \times 10^{-13} \int n_e B \cos \theta ds \quad (\text{rad m}^{-2}) \quad (9)$$

with  $B$  measured in Gauss and electron density in  $\text{cm}^{-3}$ . FORWARD carries out this integral and will report either  $FR$  at a specific frequency or  $RM$ .

A related and useful quantity that is obtained in conjunction with Faraday rotation measurements of pulsars is dispersion measure,

$$DM = \int n_e ds \quad (10)$$

which in radio astronomy is usually measured in units of  $\text{cm}^{-3}$  parsecs. The dispersion measure is used to remove the frequency-dependent delay of pulsar pulses introduced by the variation of refractive index in the interstellar and interplanetary media with



frequency, so that pulses can be aligned across the full observing bandwidth, a necessary step in measuring the rotation of the plane of polarization vs. frequency. FORWARD provides the same quantity but referred to as a column density (accessed via keyword `/colden` or in the Physical Diagnostics drop-down menu of the widget, in units of  $\text{cm}^{-2}$ ). Variations in  $DM$  on timescales of tens of minutes to hours are dominated by density variations in the solar wind, and knowledge of  $DM$  is valuable when trying to assess the relative roles of density and magnetic field in observed  $RM$  variability.

#### 4. OBSERVATIONS

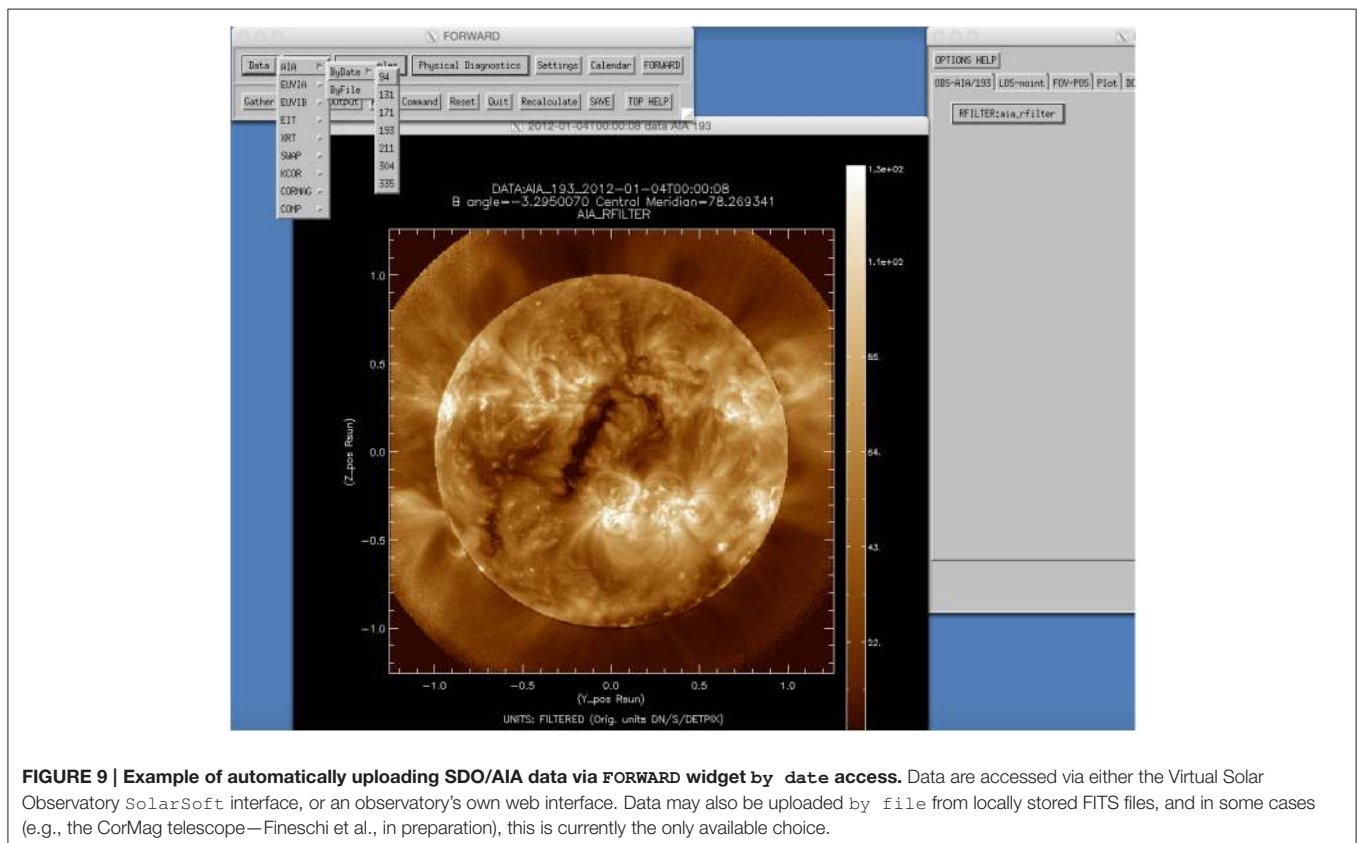
As described above in Section 3 and summarized in **Table 1**, multiple physical processes operating in the corona have sensitivities to the coronal magnetic field and manifest observable signatures from radio to soft-X-ray emission. There is thus clear value in obtaining observations at a broad range of wavelengths for intercomparison and use in constraining and defining models.

To facilitate model-data comparisons, FORWARD enables the access and manipulation of observations in a form designed to match the output of forward calculations. To this end, FORWARD extracts SolarSoft IDL maps from FITS-format observational data files, and preserves these along with associated structures in a standard format. Observational data are accessed either “by date” or “by file”—via keywords if using the command-line version of FORWARD, or a calendar/directory search if

using the widget interface (see **Figure 9**). Note that FORWARD looks first in a user-defined working directory for existing FORWARD-formatted maps or fits files before downloading or processing new files.

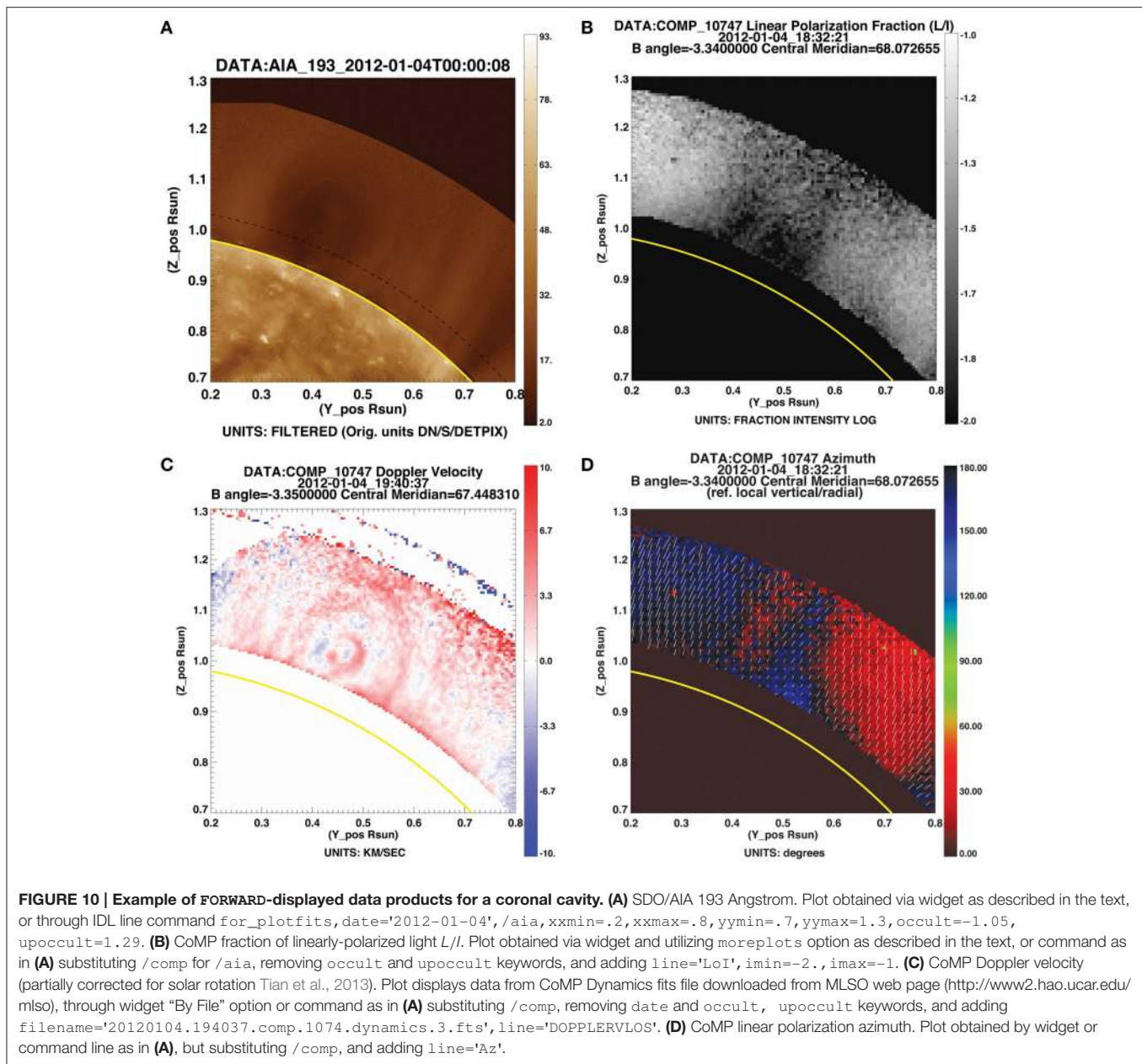
Since the focus of this paper is magnetometry, we first describe how FORWARD enables access and manipulation of Stokes polarimetric data. The Mauna Loa Solar Observatory (MLSO) Coronal Multi-channel Polarimeter (CoMP) on Hawaii (Tomczyk et al., 2008) is a 20-cm aperture coronagraph with a full field of view of the corona from 1.05 to 1.38 solar radii. It utilizes a narrow-band imaging polarimeter to observe the Fe XIII coronal line at 1074.7 and 1079.8 nm and the chromospheric He I line at 1083 nm. Data products currently served include intensity, Doppler velocity, line width, and Stokes linear polarization [Q, U as well as  $L = (Q^2 + U^2)^{1/2}$  and  $Az = 0.5 * \text{atan}(\frac{U}{Q})$ ]. CoMP linear polarization is currently the only direct magnetic diagnostic of the corona publicly available on a near-daily basis (subject to weather, etc.). These data are available online, beginning from May 2011, and can be downloaded in FITS or image format via the MLSO web pages (<http://www2.hao.ucar.edu/mlso>).

FORWARD offers another means of downloading CoMP data, and moreover acts as a tool for its display and analysis (see Gibson, 2015b for further details). **Figures 10B,D** illustrate FORWARD linear polarization output given a specified calendar date and field of view. In this case, CoMP standard “Quick Invert” data file for that date is automatically accessed, which represents an averaged image and may not include all CoMP data products.



**FIGURE 9 | Example of automatically uploading SDO/AIA data via FORWARD widget by date access.** Data are accessed via either the Virtual Solar Observatory SolarSoft interface, or an observatory’s own web interface. Data may also be uploaded by file from locally stored FITS files, and in some cases (e.g., the CorMag telescope—Fineschi et al., in preparation), this is currently the only available choice.





Comprehensive, non-averaged data are available in the “Daily Dynamics” and “Daily Polarization” FITS archives on the MLSO web pages, and once these are downloaded to a local directory they may be displayed “by file” using FORWARD. This has been done to show the CoMP Doppler velocity image of **Figure 10C**. White light data from the MLSO K-coronagraph (KCOR) can similarly be downloaded and displayed through the FORWARD widget tools.

A range of extreme ultraviolet (EUV) and soft X-ray (SXR) imager data is available through the Virtual Solar Observatory (VSO; Hill et al., 2009) and accessed by FORWARD. These include data from the currently operating Solar Dynamics Observatory Atmospheric Imaging Assembly (SDO/AIA) and Hinode X-ray Telescope (XRT), along with prior data from the

Solar and Heliospheric Observatory Extreme ultraviolet Imaging Telescope (SOHO/EIT) and Transition Region and Coronal Explorer (TRACE), which provide observations at wavelengths spanning coronal and transition region temperatures. Data from the Proba-2 “Sun Watcher using APS and Image Processing” (SWAP) EUV imager (Halain et al., 2013; Seaton et al., 2013) provide an extended (54 arcminute) field of view (FOV), and the Solar-Terrestrial Relations Observatory Extreme Ultraviolet Imagers (STEREO/EUVIA and EUVIB) provide additional viewing options for EUV coronal structures.

It is particularly simple to intercompare observations using the widget interface. For example, AIA data may be loaded by date as in **Figure 9**, and then a particular structure such as the cavity shown in **Figure 10A** may be zoomed in on using the field

of view (FOV-POS) tab of the right-hand widget. By switching on the keyword `moreplots` (located in the Output tab of the top-left widget), the zoomed-in FOV is retained and closest date/time sought in subsequent loading of CoMP data (b-d). This capability extends to forward-modeled synthetic data: if the `moreplots` option is turned on and followed by choice of a model and a click on the FORWARD button (top left widget), the field of view, viewer's position, and observable (instrument and line) are all preserved in subsequent forward calculations unless explicitly changed. In this manner, as demonstrated by Figures 11, 12 and discussed in the next section, model predictions may be directly compared to observations.

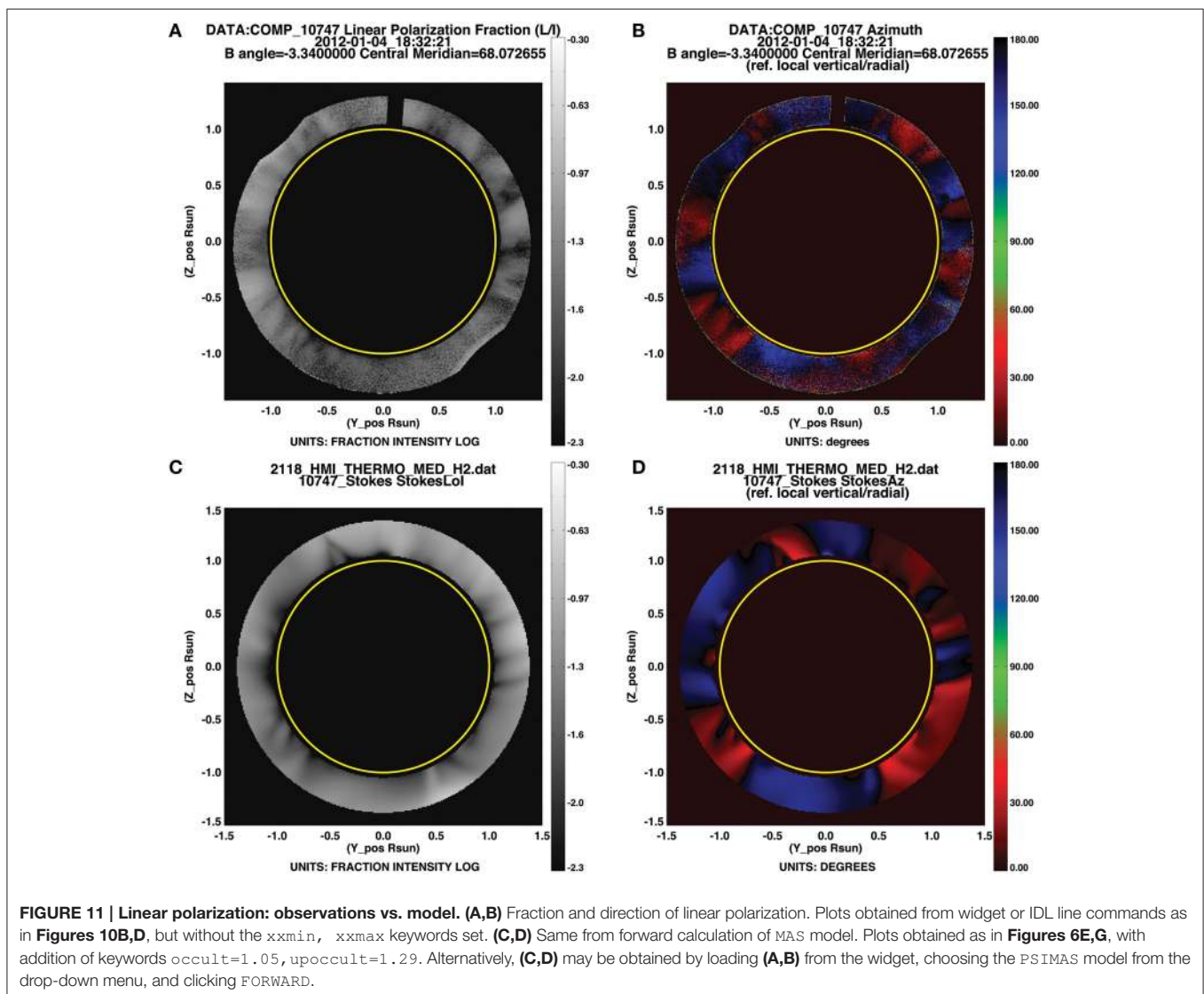
## 5. MULTIWAVELENGTH MAGNETOMETRY

Having described how FORWARD incorporates the three essential components of physical state, physical process, and observation, we now discuss how it may be used to further

multiwavelength magnetometry. FORWARD can be applied to validating models, to building intuition into how magnetic fields manifest in observations, to forward-fitting models to data, and ultimately, to developing coronal magnetic inversion methods that take full advantage of multiwavelength observations.

### 5.1. Comparing Models and Data

Figure 11 shows the CoMP observations for the day profiled in most of the Figures so far, allowing validation of the predictions of the MAS model. Inspection shows that while some regions match very well, others do not. For example, Figure 12 illustrates that for the southeast quadrant, the open region of diverging field (just below the equator) is well-captured by the model. Some of the details south of this are not captured, but the red-black-blue interface characteristic of a large-scale closed structure is reproduced. In contrast, the northwest quadrant shows considerably more structure in linear polarization observations (Figure 11A) than in the model (Figure 11C). Figure 10 further



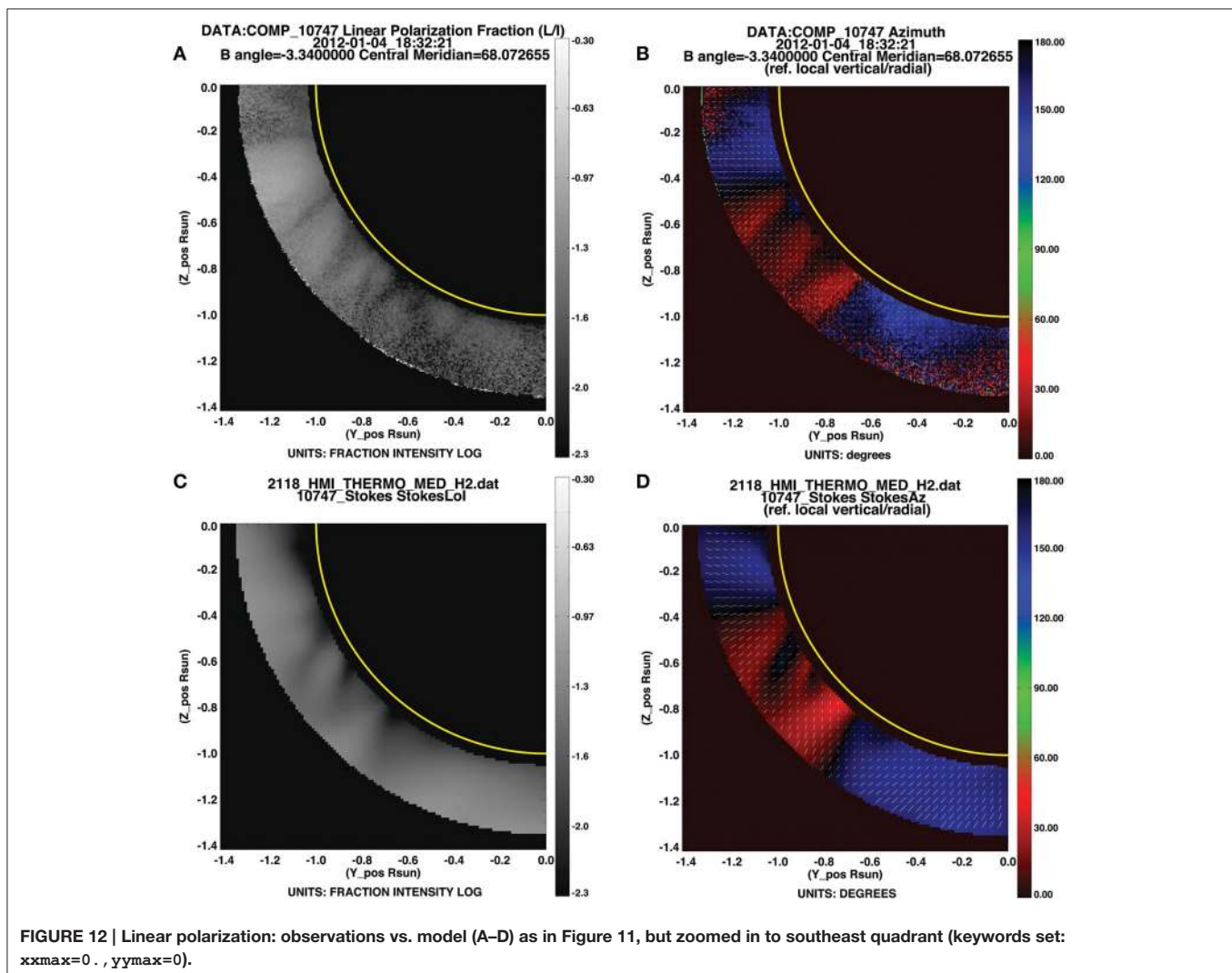
**FIGURE 11 | Linear polarization: observations vs. model.** (A,B) Fraction and direction of linear polarization. Plots obtained from widget or IDL line commands as in Figures 10B,D, but without the `xxmin`, `xxmax` keywords set. (C,D) Same from forward calculation of MAS model. Plots obtained as in Figures 6E,G, with addition of keywords `occult=1.05`, `upocult=1.29`. Alternatively, (C,D) may be obtained by loading (A,B) from the widget, choosing the PSIMAS model from the drop-down menu, and clicking FORWARD.

magnifies the data in this region, and demonstrates that the northern-most of these linear polarization structures is associated with a coronal cavity. The difference between model and data in this region likely arises because, although the MAS simulation we have shown is non-potential, it does not capture all coronal currents. In particular, currents that slowly build up over time are not reproduced. Such a buildup of currents is expected in polar crown regions (Yeates and Mackay, 2012), and is thus likely in the region of the cavity of **Figure 10**.

The linear polarization observations of the CoMP telescope represent a unique observational resource, and one that has benefited greatly from the intuition built via forward modeling. In advance of CoMP's synoptic operation at MLSO, the CLE code was used to demonstrate that the presence of currents in the corona should be observable in Fe XIII linear polarization (Judge et al., 2006). Indeed, this has proved the case, and comparisons of CoMP data to FORWARD-generated images have shown linear polarization to be a useful diagnostic of magnetic topologies, including spheromaks (Dove et al., 2011),

pseudostreamers (Rachmeler et al., 2014), and cylindrical flux ropes (Bağ-Stęślicka et al., 2013; Rachmeler et al., 2013). The CoMP linear polarization structure shown in **Figure 10B** is an example of a “lagomorph” (named for its rabbit's-head shape). Bağ-Stęślicka et al. (2013, 2014) studied dozens of examples of CoMP lagomorphs, and showed clear correlation with the size and location of associated EUV prominence cavities (e.g., **Figure 10A** vs. **Figure 10B**). The authors also used forward modeling to demonstrate that a magnetic flux rope model results in a lagomorph: the van Vleck angles within the outer portions of the flux rope and the overlying arcade creates a dark structure framing the rabbit's ears and the sides of its head, and sheared or twisted fields at the flux rope's axis, being oriented perpendicular to the plane of sky, are also relatively dark in linear polarization and form the center of the head.

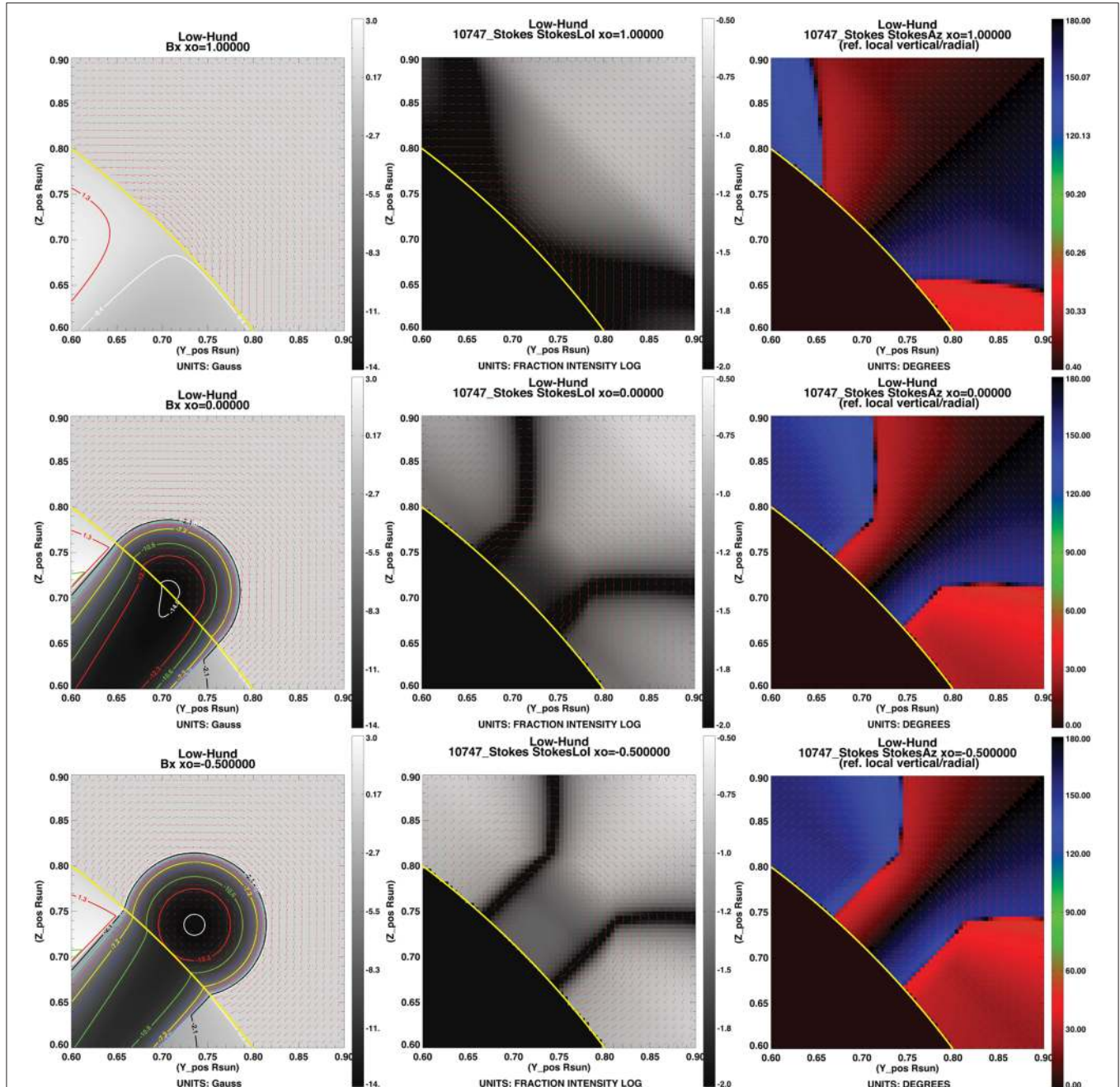
The presence of a wide V shape is expected even for a potential field arcade. Using an analytic model of a flux rope included in FORWARD (Low and Hundhausen, 1995), **Figure 13** shows how the addition of coronal currents above the magnetic neutral





line narrows this V and introduces a dark central structure. The bottom row of **Figure 13** may be compared to **Figure 10**, noting that the top of the cavity/flux rope is near the top of the CoMP field of view, so that the ears are not captured in this case.

As we have discussed above, there are a range of multiwavelength observations that can be used to constrain coronal magnetic fields. Indeed, coronal-cavity white light and emission observations have been interpreted as largely independent indicators of a flux-rope magnetic structure (see



**FIGURE 13 |** Using an analytic flux-rope model (Low and Hundhausen, 1995), we see that the presence of currents above the underlying neutral line narrows the ears and introduces the dark central head structure to a linear polarization lagomorph. Left column: LOS-oriented magnetic field strength. Plots obtained through commands `for_drive, 'lowhund', line='bx', thetao=45., x_oinput=xo, xxmin=0.6, xxmax=0.9, yymin=0.6, ymax=0.9, /fieldlines, imin=-14, imax=3` for values of  $x_o = 1., 0., -.5$ . Middle column: LOS-integrated linear polarization fraction. Plots obtained from same commands, substituting `line='LoI', imax=-.5, imin=-2.` and adding keyword `/comp`. Right column: LOS-integrated linear polarization direction (azimuth). Plots obtained as for `L/I`, but substituting `line='Az'`.



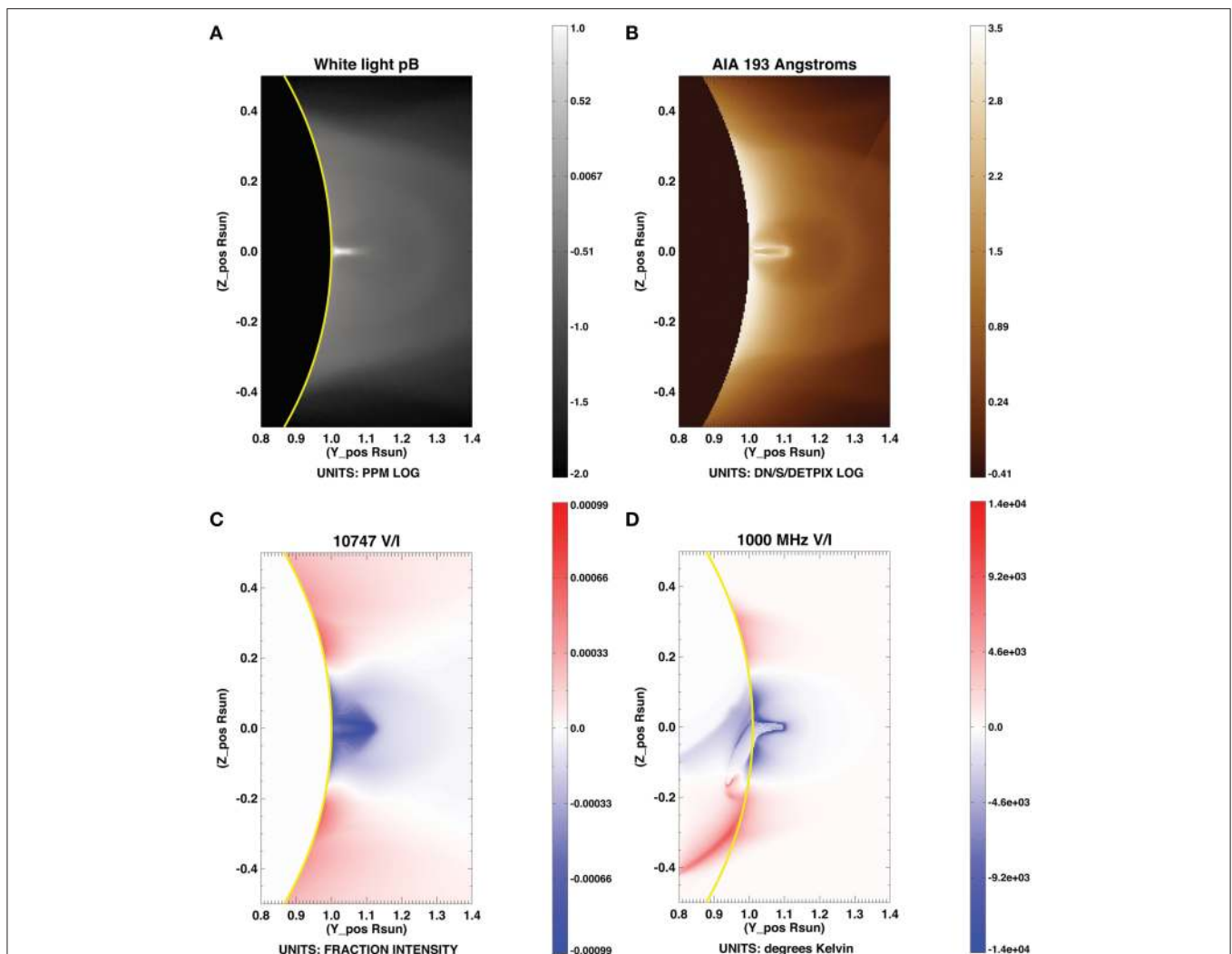
discussion in Gibson, 2014, 2015a; see also Bąk-Stęślicka et al., in preparation for discussion of LOS flows as indications of magnetic-flux-rope topology). Quantification of the three-dimensional morphology, substructure, and plasma properties of cavities have served to justify such interpretations. These quantifications were obtained by fitting the “CAVMORPH” analytic model included within the FORWARD distribution to observations of cavities in white light, EUV, and SXR (Gibson et al., 2010; Schmit and Gibson, 2011; Kucera et al., 2012; Reeves et al., 2012).

Such “forward fitting” goes beyond intuition building, and in fact is a means of inverting observations to quantify properties of the physical state. It does require specification of a parameterized model, such that through iteration best-fit parameters are determined. Dalmasse et al. (in preparation) provides an example

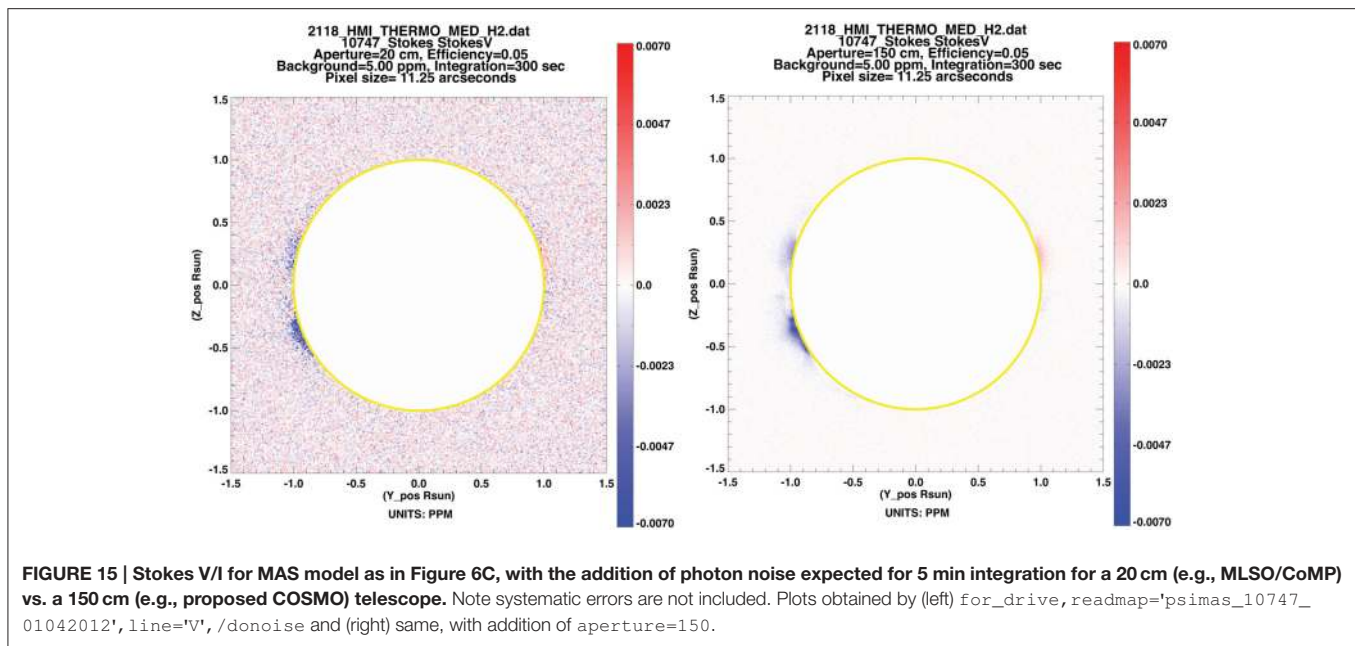
of a statistical method applied to forward fitting a flux-rope model to visible/IR polarimetric data, including both linear and circular polarization (see also Jibben et al., submitted). Other inversion methods applicable to these data are also under development (Kramar et al., 2006, 2013, 2014; Plowman, 2014; also Kramar et al., in preparation).

## 5.2. Synthetic Testbeds and Beyond

The method described in Dalmasse et al. (in preparation) employs synthetic Fe XIII linear polarization data generated using FORWARD. This work represents an area of active development for FORWARD, i.e., the creation of multiwavelength synthetic data for coronal magnetic structures ranging from active regions (e.g., M. Rempel, private communication), to polar crown prominence/cavity systems



**FIGURE 14 | Synthetic data including (A) visible and (B) EUV intensity, and (C) IR and (D) radio circular polarization, generated for simulated prominence-cavity system (Fan, personal communication). Plot (A) obtained by FORWARD line command: `for_drive, 'numcube', cubename='$FORWARD_DB/TESTBEDS/fullthermodynamic_erupting_qp_mhd_Fan', xxmin=0.8, xxmax=1.4, yymin=-.5, ymax=.5, units='PPM', cuberot=-25., colortable=0`. Plot (B) obtained with addition of keyword `/aia`. Plot (C) added with addition of keyword `/comp, line='VOI'`, and (D) with addition of `/radio, line='VOI'`.**



(e.g., Fan, personal communication; see **Figure 14**, to a global corona containing a variety of currents, e.g., D. Mackay, private communication). These numerical simulations will be included in future FORWARD SolarSoft distributions, and from them synthetic data ranging from radio to SXR wavelengths can be generated as community testbeds to aid in the development of inversion methods and in analyses of the sensitivity of different types of observations to physical parameters.

The images shown in **Figures 1–14** are idealized. Real data has noise, and inversions must take this into account. Sensitivity, field of view and spatial resolution of the telescope used to obtain the data may all contribute to noise. For polarimetry, sensitivity is usually a constraining factor because (at least at optical, IR and UV wavelengths) the polarized signals are much weaker than the total intensities and subject to cross-talk that requires careful time-consuming calibration to ensure robust measurements. With modern telescopes one is also often trading large fields of view for high spatial resolution, and for the observation of spatially extended features this can be a problem.

At radio wavelengths, one has to deal with the fact that spatial resolution is always frequency dependent: for a fixed effective aperture dimension, the size of a resolution element is inversely proportional to frequency, and for typical modern radio observations taken over a wide frequency range, the spatial resolution can vary by factors of several from high to low frequencies. For the purpose of measuring coronal magnetic fields with gyroresonance emission where field strength is proportional to frequency, this means that one generally has poorer spatial resolution for the study of weak fields than for strong fields. Fortunately this is in the right direction since strong-field regions are usually smaller, but it does limit our ability to study three-dimensional fields with uniform resolution.

At the moment, FORWARD only implements noise for visible/IR spectropolarimetry (Section 3.4), and then only photon

noise (see **Figure 15**). Efforts are underway to allow “instrument personality profiles” in which a loss of resolution appropriate to a particular observation could be overlaid on the forward calculation, given details of a telescope and its observing configuration. This capability would enable the design and use of future large telescopes such as the Daniel K. Inouye Solar Telescope (DKIST), the Frequency Agile Solar Telescope (FASR), and the Coronal Solar Magnetism Observatory (COSMO; e.g., **Figure 15**; see also Lin, 2016).

## 6. CONCLUSIONS

Our primary motivation in developing FORWARD has been to enable multiwavelength coronal magnetometry. The coronal magnetic field lies at the heart of many of the mysteries of solar physics, including coronal heating, solar wind acceleration, and flare and coronal mass ejection onset and evolution. It holds the key to progress in predictive capability for space-weather events: in particular, the direction of the magnetic field at 1 AU depends crucially on the magnetic field at its coronal source, and on the context of this source in both time and space. In this paper we have demonstrated how different physical processes effectively highlight different aspects of the coronal magnetic field, and how these manifest in observations at different wavelengths. Because the photospheric magnetic field is not force-free, our ability to find a meaningful solution to coronal magnetic field through extrapolations from this boundary is limited (De Rosa et al., 2009). We therefore *must* make use of multiwavelength observations of the solar atmosphere to further constrain the global coronal magnetic field.

FORWARD represents a community effort to design and gather a library of codes for the synthesis of multiwavelength coronal data from physical models. Our philosophy has been

to incorporate as many existing resources as possible, and to make use of the comprehensive and ever-growing resources available via SolarSoft IDL. We note complementary capabilities available for forward modeling in radio wavelengths, i.e., the *GX\_Simulator* package referred to above in Section 3.6 (Nita et al., 2015), and for forward modeling coronal waves, i.e., the *FoMo* codes described in van Doorselaere et al. (2016).

FORWARD continues to be developed. New subroutines for ultraviolet spectropolarimetry in the unsaturated Hanle regime are being tested (Fineschi, 2001; see also Raouafi et al., submitted; Dima et al., submitted). We are also expanding our numerical interface to allow varied-grid models (currently numerical datacubes must be on a regular grid). A future goal will be to add capability for synthesizing heliospheric images, which would complement current capability for Faraday rotation, and enable connections between imaging and *in situ* observations during the era of Solar Probe and Solar Orbiter. The wide variety of multiwavelength data currently and soon to be available, in combination with ongoing efforts to develop comprehensive and efficient inversion methods, makes us confident that ultimately the goal of quantifying the coronal magnetic field will be achieved.

## AUTHOR CONTRIBUTIONS

SG was primary author of this paper, wrote or contributed to most of the subroutines in FORWARD, and is responsible for the oversight of its ongoing development. TK wrote Sections 3.2, 3.3, and contributed to Sections 5, 6. She wrote or otherwise coordinated all FORWARD codes and databases related to EUV/SXR imaging and spectroscopy, and contributed to several other FORWARD subroutines. SW wrote Sections 3.6, 3.7, and contributed to Section 5, 6. He wrote or otherwise coordinated all FORWARD codes related to radio spectropolarimetry. JD was the main author of the integration of CLE into FORWARD, and the initial developer of several of the backbone subroutines of FORWARD as well as the analytic model interface. LR developed the numerical interface and incorporated the PFSS model into FORWARD, and contributed to the data-download interface codes and other FORWARD subroutines. YF contributed to the numerical interface effort, provided an axisymmetric flux rope model to the FORWARD package for demonstration purposes, and is providing numerical simulations for synthetic test-beds as discussed in Section 5. BF was the initial developer of the *for\_widget* interface. CD contributed a subroutine for interfacing with the PSI MAS model and assisted with its

implementation. KR contributed the XRT response subroutine. All authors read and critically revised the paper, approved the final version, and agreed to be accountable for all aspects of the work.

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