

# Gas entropy in a representative sample of nearby X-ray galaxy clusters (REXCESS): relationship to gas mass fraction

G.W. Pratt<sup>1,2</sup>, M. Arnaud<sup>1</sup>, R. Piffaretti<sup>1</sup>, H. Böhringer<sup>2</sup>, T.J. Ponman<sup>3</sup>, J.H. Croston<sup>4</sup>, G.M. Voit<sup>5</sup>, S. Borgani<sup>6</sup>  
and R.G. Bower<sup>7</sup>

<sup>1</sup> Laboratoire AIM, IRFU/Service d'Astrophysique - CEA/DSM - CNRS - Université Paris Diderot, Bât. 709, CEA-Saclay, F-91191 Gif-sur-Yvette Cedex, France

<sup>2</sup> Max-Planck-Institut für extraterrestrische Physik, Giessenbachstraße, 85748 Garching, Germany

<sup>3</sup> School of Physics and Astronomy, University of Birmingham, Edgbaston, Birmingham B15 2TT, UK

<sup>4</sup> School of Physics and Astronomy, University of Southampton, Southampton, Hampshire, SO17 1BJ, UK

<sup>5</sup> Department of Physics and Astronomy, Michigan State University, East Lansing, MI 48824-2320, USA

<sup>6</sup> Dipartimento di Astronomia dell'Università di Trieste, via Tiepolo 11, 34131 Trieste, Italy

<sup>7</sup> Institute for Computational Cosmology, Department of Physics, Durham University, South Road, Durham DH1 3LE, UK

Received 17 September 2009; accepted 10 November 2009

## ABSTRACT

We examine the radial entropy distribution and its scaling using 31 nearby galaxy clusters from the Representative *XMM-Newton* Cluster Structure Survey (REXCESS), a sample in the temperature range 2-9 keV selected in X-ray luminosity only, with no bias toward any particular morphological type. The entropy profiles are robustly measured at least out to  $R_{1000}$  in all systems and out to  $R_{500}$  in thirteen systems. Compared to theoretical expectations from non-radiative cosmological simulations, the observed distributions show a radial and mass-dependent excess entropy, such that the excess is greater and extends to larger radii in lower mass systems. At  $R_{500}$ , the mass dependence and entropy excess are both negligible within the large observational and theoretical uncertainties. Mirroring this behaviour, the scaling of gas entropy is shallower than self-similar in the inner regions, but steepens with radius, becoming consistent with self-similar at  $R_{500}$ . There is a large dispersion in scaled entropy in the inner regions, apparently linked to the presence of cool cores and dynamical activity; at larger radii the dispersion decreases by approximately a factor of two to 30 per cent, and the dichotomy between subsamples disappears. There are two peaks in the distribution of both inner slope and, after parameterising the profiles with a power law plus constant model, in central entropy  $K_0$ . However, we are unable to distinguish between a bimodal or a left-skewed distribution of  $K_0$  with the present data. The distribution of outer slopes is unimodal with a median value of 0.98, and there is a clear correlation of outer slope with temperature. Renormalising the dimensionless entropy profiles by the gas mass fraction profile  $f_{\text{gas}}(< R)$ , leads to a remarkable reduction in the scatter, implying that gas mass fraction variations with radius and mass are the cause of the observed entropy structural and scaling properties. The results are consistent with the picture of a cluster population in which entropy modification is centrally concentrated and extends to larger radii at lower mass, leading to both a radial and a mass-dependence in the gas mass fraction, but which is increasingly self-similar at large radius. The observed normalisation, however, would suggest entropy modification at least up to  $R_{1000}$ , and even beyond, in all but the most massive systems. We discuss a tentative scenario to explain the observed behaviour of the entropy and gas mass fraction in the REXCESS sample, in which a combination of extra heating and merger mixing maintains an elevated central entropy level in the majority of the population, and a smaller fraction of systems is able to develop a cool core.

**Key words.** Cosmology: observations, Cosmology: dark matter, Galaxies: cluster: general, (Galaxies) Intergalactic medium, X-rays: galaxies: clusters

## 1. Introduction

The first order model of structure formation – that of hierarchical, dark matter dominated gravitational collapse – is capable of reproducing only the gross statistical properties of the galaxy cluster population. In this scenario, the intracluster medium (ICM) is heated to X-ray emitting temperatures by shocks and compression as it falls into the potential well of the dark matter, and the resulting X-ray cluster population is self-similar and scale-free. In real clusters, second order effects, linked primarily to feedback from galaxy formation and radiative cooling of the gas, serve to modify the X-ray properties of the ICM with respect to these expectations (see, e.g., Voit 2005; Borgani & Kravtsov 2009, for recent reviews). The effect of these nongravitational

processes is substantial. It can be seen most readily in the relation between the X-ray luminosity and temperature, which in the first order scenario scales simply as  $L \propto T^3$ , but which is observed to scale as  $L \propto T^3$  (e.g. Pratt et al. 2009, and references therein), implying a progressive suppression of luminosity in low temperature systems.

In recent years, spatially resolved observations have allowed us to examine in more detail the impact of nongravitational processes on the ICM, mainly through radial profiles and mapping. In this context, the entropy  $K$  of the ICM<sup>1</sup> is of considerable interest because the observable X-ray characteristics of a cluster

<sup>1</sup> Keeping with convention, we use the X-ray astronomer's 'entropy' throughout this paper. Defined as  $K = kT/n_e^{2/3}$ , where  $n_e$  is the electron number density, this quantity is related to the true entropy by a logarithm and an additive constant.

**Table 1.** Basic cluster data.  $K(R_s)$  is the entropy measured at the radius corresponding to density contrast  $\delta$ . Cool core and morphological classification from Pratt et al. (2009).

Cluster	$z$	$kT$	$M_{500}$	$K(0.1 R_{200})$	$K(R_{2500})$	$K(R_{1000})$	$K(R_{500})$	CC	Disturbed
RXC J0003.8+0203	0.0924	$3.83^{+0.11}_{-0.11}$	$2.11^{+0.04}_{-0.04}$	$273.35 \pm 17.67$	$619.88 \pm 50.40$	$913.29 \pm 114.31$	...	...	...
RXC J0006.0-3443	0.1147	$5.24^{+0.23}_{-0.21}$	$3.95^{+0.12}_{-0.12}$	$469.36 \pm 55.82$	$786.71 \pm 80.54$	$1040.90 \pm 98.14$	...	...	✓
RXC J0020.7-2542	0.1410	$5.54^{+0.13}_{-0.13}$	$3.84^{+0.06}_{-0.06}$	$347.10 \pm 28.49$	$727.97 \pm 58.46$	$1085.23 \pm 143.63$	$1498.66 \pm 687.71$	...	...
RXC J0049.4-2931	0.1084	$2.87^{+0.10}_{-0.10}$	$1.62^{+0.04}_{-0.04}$	$191.06 \pm 14.02$	$442.17 \pm 74.37$	$647.11 \pm 82.13$	...	...	...
RXC J0145.0-5300	0.1168	$5.81^{+0.15}_{-0.15}$	$4.37^{+0.08}_{-0.08}$	$365.72 \pm 26.31$	$840.65 \pm 58.67$	$1400.17 \pm 166.17$	...	...	✓
RXC J0211.4-4017	0.1008	$2.08^{+0.05}_{-0.05}$	$1.00^{+0.02}_{-0.02}$	$141.17 \pm 8.83$	$309.39 \pm 22.19$	$522.67 \pm 47.23$	$687.20 \pm 134.45$	...	...
RXC J0225.1-2928	0.0604	$2.53^{+0.14}_{-0.14}$	$0.96^{+0.04}_{-0.04}$	$173.95 \pm 13.33$	$546.12 \pm 114.66$	$747.10 \pm 81.97$	...	...	✓
RXC J0345.7-4112	0.0603	$2.28^{+0.07}_{-0.07}$	$0.97^{+0.02}_{-0.02}$	$175.30 \pm 13.91$	$384.26 \pm 38.58$	$481.23 \pm 40.90$	...	✓	...
RXC J0547.6-3152	0.1483	$6.04^{+0.14}_{-0.14}$	$4.98^{+0.09}_{-0.09}$	$324.96 \pm 16.61$	$842.16 \pm 61.88$	$1128.97 \pm 96.51$	...	...	...
RXC J0605.8-3518	0.1392	$4.93^{+0.12}_{-0.12}$	$3.87^{+0.06}_{-0.06}$	$238.61 \pm 13.12$	$631.88 \pm 32.69$	$1167.71 \pm 157.17$	$1745.33 \pm 462.31$	✓	...
RXC J0616.8-4748	0.1164	$4.18^{+0.11}_{-0.11}$	$2.70^{+0.06}_{-0.06}$	$348.61 \pm 39.07$	$639.08 \pm 40.11$	$939.62 \pm 77.66$	$1357.27 \pm 224.37$	...	✓
RXC J0645.4-5413	0.1644	$7.23^{+0.18}_{-0.18}$	$7.38^{+0.14}_{-0.14}$	$349.48 \pm 27.52$	$941.30 \pm 79.96$	$1462.99 \pm 158.42$	...	...	...
RXC J0821.8+0112	0.0822	$2.81^{+0.10}_{-0.10}$	$1.31^{+0.03}_{-0.03}$	$285.87 \pm 19.08$	$436.36 \pm 45.57$	$617.55 \pm 56.35$	...	...	...
RXC J0958.3-1103	0.1669	$5.95^{+0.19}_{-0.19}$	$4.17^{+0.15}_{-0.15}$	$220.23 \pm 33.22$	$875.57 \pm 297.68$	$1421.19 \pm 618.50$	...	✓	...
RXC J1044.5-0704	0.1342	$3.58^{+0.05}_{-0.05}$	$2.27^{+0.02}_{-0.02}$	$164.67 \pm 6.09$	$447.65 \pm 21.65$	$722.02 \pm 84.78$	$1021.75 \pm 127.41$	✓	...
RXC J1141.4-1216	0.1195	$3.58^{+0.06}_{-0.06}$	$2.27^{+0.02}_{-0.02}$	$197.66 \pm 7.26$	$557.71 \pm 25.24$	$849.48 \pm 74.07$	$1016.40 \pm 91.66$	✓	...
RXC J1236.7-3354	0.0796	$2.77^{+0.06}_{-0.06}$	$1.33^{+0.02}_{-0.02}$	$179.13 \pm 6.56$	$469.68 \pm 25.36$	$849.80 \pm 133.02$	...	...	...
RXC J1302.8-0230	0.0847	$3.48^{+0.08}_{-0.08}$	$1.89^{+0.03}_{-0.03}$	$310.02 \pm 16.16$	$741.95 \pm 43.99$	$779.02 \pm 89.95$	$684.30 \pm 52.30$	✓	✓
RXC J1311.4-0120	0.1832	$8.67^{+0.12}_{-0.12}$	$8.41^{+0.08}_{-0.08}$	$337.47 \pm 10.26$	$1044.38 \pm 56.40$	$1653.59 \pm 116.17$	...	✓	...
RXC J1516.3+0005	0.1181	$4.68^{+0.13}_{-0.13}$	$3.28^{+0.07}_{-0.07}$	$321.83 \pm 16.45$	$673.65 \pm 35.14$	$1004.69 \pm 96.87$	...	...	...
RXC J1516.5-0056	0.1198	$3.70^{+0.10}_{-0.10}$	$2.59^{+0.03}_{-0.03}$	$314.75 \pm 22.64$	$712.59 \pm 51.51$	$728.24 \pm 54.37$	$732.28 \pm 50.80$	...	✓
RXC J2014.8-2430	0.1538	$5.75^{+0.10}_{-0.10}$	$5.38^{+0.07}_{-0.07}$	$263.67 \pm 12.80$	$726.64 \pm 39.70$	$1158.45 \pm 96.11$	...	✓	...
RXC J2023.0-2056	0.0564	$2.72^{+0.09}_{-0.09}$	$1.21^{+0.03}_{-0.03}$	$248.64 \pm 15.10$	$527.38 \pm 58.61$	$657.84 \pm 42.38$	...	...	✓
RXC J2048.1-1750	0.1475	$5.06^{+0.11}_{-0.11}$	$4.32^{+0.07}_{-0.07}$	$409.74 \pm 35.68$	$743.70 \pm 46.76$	$1006.16 \pm 53.29$	$1086.97 \pm 79.89$	...	✓
RXC J2129.8-5048	0.0796	$3.84^{+0.14}_{-0.14}$	$2.26^{+0.06}_{-0.06}$	$441.52 \pm 31.94$	$644.40 \pm 53.66$	$861.48 \pm 67.50$	...	...	✓
RXC J2149.1-3041	0.1184	$3.48^{+0.07}_{-0.07}$	$2.25^{+0.03}_{-0.03}$	$196.26 \pm 9.85$	$513.99 \pm 32.71$	$762.82 \pm 80.24$	$1062.45 \pm 185.45$	✓	...
RXC J2157.4-0747	0.0579	$2.79^{+0.07}_{-0.07}$	$1.29^{+0.03}_{-0.03}$	$390.50 \pm 40.35$	$659.44 \pm 65.91$	$675.83 \pm 23.62$	...	...	✓
RXC J2217.7-3543	0.1486	$4.62^{+0.09}_{-0.09}$	$3.61^{+0.05}_{-0.05}$	$282.04 \pm 13.59$	$610.69 \pm 34.87$	$872.97 \pm 62.41$	$1076.81 \pm 107.72$	...	...
RXC J2218.6-3853	0.1411	$6.18^{+0.20}_{-0.20}$	$4.92^{+0.11}_{-0.11}$	$263.03 \pm 12.24$	$859.14 \pm 64.35$	$1610.20 \pm 225.47$	$2278.44 \pm 862.21$	...	✓
RXC J2234.5-3744	0.1510	$7.32^{+0.12}_{-0.12}$	$7.36^{+0.09}_{-0.09}$	$451.39 \pm 28.73$	$813.38 \pm 42.42$	$1170.94 \pm 97.12$	...	...	...
RXC J2319.6-7313	0.0984	$2.56^{+0.08}_{-0.08}$	$1.56^{+0.03}_{-0.03}$	$133.95 \pm 7.23$	$380.37 \pm 32.90$	$637.12 \pm 66.10$	$1017.21 \pm 211.39$	✓	✓

<sup>a</sup> Temperature in keV, measured in the  $[0.15 - 0.75] R_{500}$  aperture

<sup>b</sup> Mass in units of  $h_{70}^{-1} 10^{14} M_{\odot}$ , estimated from the  $M_{500} - Y_X$  relation given in Eqn. 1 (Arnaud et al. 2009).

are just manifestations of its distribution in the dark matter potential well. Entropy is generated during the hierarchical assembly process, yet is modified by any other process that can change the physical characteristics of the gas. It is thus a quantity that preserves a record both of the accretion history of a cluster and of the influence of non-gravitational processes on the properties of its ICM, and as such it is a useful tool for our understanding of the thermodynamic history of the cluster population.

Early measurements of the entropy based primarily on *ROSAT* and *ASCA* data indicated that groups had flatter entropy profiles than cluster scale objects (David et al. 1996), and measurements of the entropy at  $0.1 R_{200}$  revealed an entropy-temperature ( $K - T$ ) relation that was shallower than expected (Ponman et al. 1999; Lloyd-Davies et al. 2000; Ponman et al. 2003). These data also afforded the first indications for excess entropy above that expected from gravitational collapse even at large radius in group-scale objects (Finoguenov et al. 2002; Ponman et al. 2003). The advent of *XMM-Newton* and *Chandra* has allowed relatively high resolution spatially resolved measurement of the entropy to be obtained across a wide range of cluster and group masses (Pratt & Arnaud 2003; Piffaretti et al. 2005; Pratt et al. 2006; Morandi & Ettori 2007; Nagai et al. 2007; Zhang et al. 2008; Sanderson et al. 2009; Johnson et al. 2009). Recent results have suggested that the entropy is indeed

higher than expected from gravitational collapse at least out to  $R_{2500}$  (Pratt et al. 2006; Nagai et al. 2007; Sun et al. 2009), and perhaps further (Pratt et al. 2006; Sun et al. 2009), even up to relatively high masses. In addition, indications for excess entropy have been found at large radius in intermediate redshift groups (Jeltema et al. 2006), and the first measurements of the entropy evolution have been undertaken (Ettori et al. 2004).

In the present paper we re-investigate the entropy with REXCESS (Böhringer et al. 2007), a representative sample of 33 local ( $z < 0.2$ ) clusters drawn from the REFLEX catalogue (Böhringer et al. 2004), all of which have been observed with *XMM-Newton*. The properties of the REXCESS sample allow us to define a robust local reference for entropy structure and scaling. REXCESS was designed to be representative of any high-quality local X-ray survey, thus clusters have been selected in luminosity only, ensuring no morphological bias, in such a way as to sample the X-ray cluster luminosity function in an optimal manner. Moreover, distances were optimised so that  $R_{500}$  falls well within the *XMM-Newton* field of view, increasing the precision of measurements at large radii as compared to more nearby clusters, which often fill the field of view and for which background modelling is consequently more complicated.

In the following, we first examine the normalisation of the entropy with respect to predictions from non-radiative cosmo-

logical simulations – such ‘adiabatic’ simulations include only gravitational processes – finding a systematic entropy excess that is greater at small radii and in lower mass systems. The considerable dispersion at small radii appears linked to whether a cluster possesses a cool core or is morphologically disturbed. The mass dependence disappears at  $R_{500}$ , implying that entropy scaling is self-similar, with a normalisation that is approximately consistent with predictions. Parameterising the profiles in terms of a power law plus constant model, there are two peaks in the distribution of central entropy but there is no strong evidence that it is bimodal. The distribution of outer slopes is unimodal and the slope depends on temperature. We then link the entropy scaling and structural properties to a systematic variation in gas content with total mass and with radius. Finally, we discuss mechanisms which could bring about the observed entropy characteristics, and propose a tentative scenario to explain the observed entropy distributions.

We adopt a  $\Lambda$ CDM cosmology with  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$  (i.e.,  $h_{70} = 1$ ),  $\Omega_M = 0.3$  and  $\Omega_\Lambda = 0.7$ . All uncertainties are quoted at the 68 per cent confidence level.

## 2. Sample and analysis

### 2.1. Sample description and subsample definition

A full description of the REXCESS sample, including the *XMM-Newton* observation details, can be found in Böhringer et al. (2007), and the preliminary data analysis is described in Croston et al. (2008). Two of the objects, RCXC J0956.4-1004 and RXC J2157.4-0747 (the Abell 901/902 supercluster and a bimodal cluster, respectively), display complex morphologies which preclude their use for the present radial profile analysis. Basic cluster parameters are listed in Table 1.

On occasion in the following, we will subdivide the sample into cool core and non-cool core systems, or according to whether the clusters are morphologically relaxed or unrelaxed. These subsamples were established to cull approximately the most extreme thirty percent of the full sample in each category and are defined as in Pratt et al. (2009). Thus clusters with central density  $E(z)^{-2}n_{e,0} > 4 \times 10^{-2} \text{ cm}^{-3}$  are classified<sup>2</sup> as cool core systems (10/31), and those with centre shift parameter  $\langle w \rangle > 0.01 R_{500}$  (derived with the central regions excised) are classified as morphologically disturbed (12/31). Both the central densities  $n_{e,0}$  and centroid shift parameter  $\langle w \rangle$  are given in Haarsma et al. (2009).

### 2.2. Data analysis

Event lists were processed using version 7.0 of the *XMM-Newton* SAS. All data products were extracted from event lists that were generated, cleaned, PATTERN-selected, vignetting-corrected, and point source-removed as described in Pratt et al. (2007).

#### 2.2.1. Gas density profiles

The procedure used to calculate the gas density profiles, plus extensive analysis of their properties, is described in full in Croston et al. (2008). In brief, surface brightness profiles, centred on the peak of the X-ray emission, were extracted from 3''3 bins in the [0.3-2] keV band and deprojected and PSF corrected

using the non-parametric method introduced in Croston et al. (2006). These were converted to gas density by calculating a conversion factor in XSPEC using the temperature in the [0.15 – 1]  $R_{500}$  aperture, and subsequently corrected to take into account radial variations of temperature and abundance to give the final deprojected, PSF-corrected radial density profile.

#### 2.2.2. Temperature profiles

Projected 2D temperature profiles were derived from spectra extracted in logarithmically-spaced annular bins centred on the peak of the X-ray emission. Binning was such that the first bin was defined to have a significance of  $30\sigma$  above background, and subsequent bins were defined so as to have  $R_{\text{out}}/R_{\text{in}} = 1.33 - 1.5$  depending on the quality of the observation. The instrumental and particle background was subtracted from each annulus using custom stacked, recast data files accumulated from observations obtained with the filter wheel in the CLOSED position (FWC), renormalised using the count rate in a high energy band free of cluster emission<sup>3</sup>. After subtraction of the FWC spectra, all spectra were grouped to a minimum of 25 counts per bin.

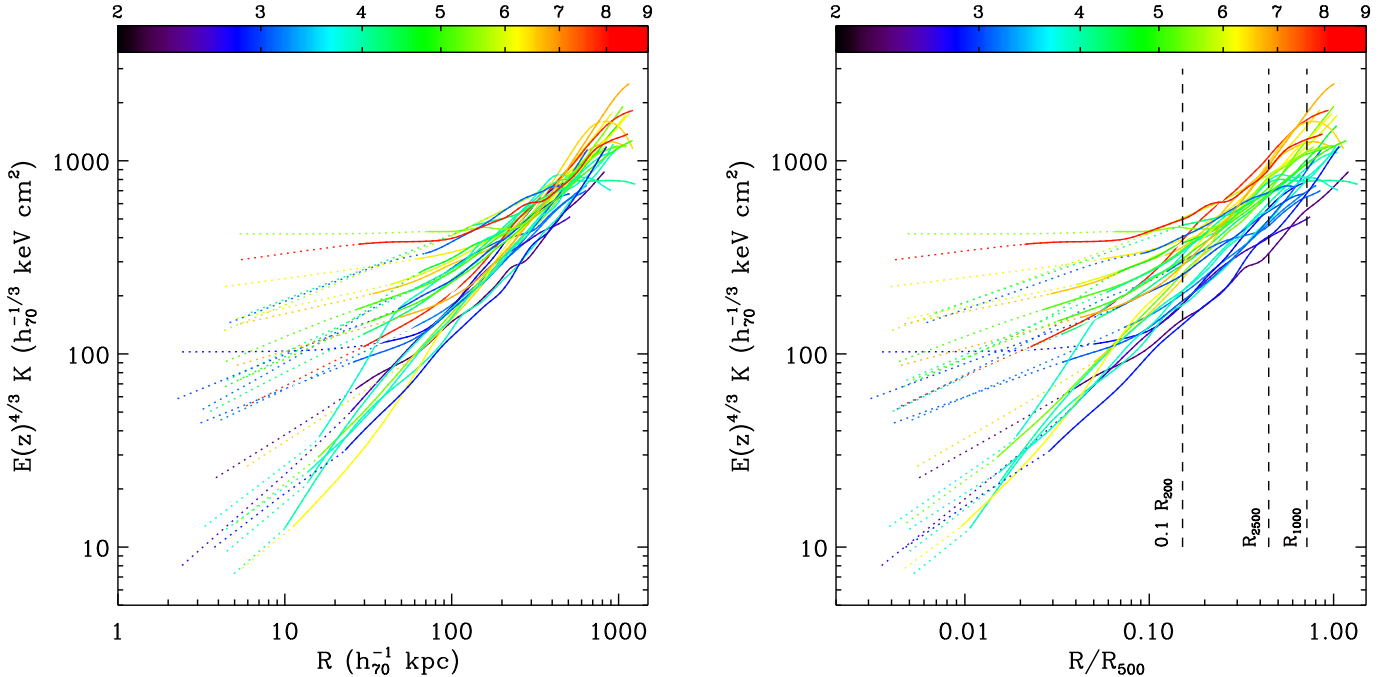
We then fitted the FWC-subtracted spectrum of an annular region external to the cluster emission with a physically motivated model of the cosmic X-ray background, consisting of two MEKAL components plus an absorbed power law with a fixed slope of  $\Gamma = 1.4$  (see Lumb et al. 2002; De Luca & Molendi 2004). Spectra were fitted in the [0.3 – 10] keV range using  $\chi^2$  statistics, excluding the [1.4-1.6] keV band (due to the Al line in all three detectors), and, in the EPN, the [7.45-9.0] keV band (due to the strong Cu line complex). In these fits the MEKAL models are unabsorbed and have solar abundances, and the temperature and normalisations are free parameters; the powerlaw component is absorbed by the Galactic absorption in the direction of the cluster and since it has a fixed slope, only its normalisation is an additional free parameter in the fit. This best fitting background model, with renormalisation appropriate to the ratio of the surface area of the extraction regions (corrected for chip gaps, etc.), is then added as an extra component in each annular fit. This is our standard fit.

We then vary the power law index in such a way as to mimic a  $\pm 10$  per cent variation in the [2-10] keV flux and refit the spectrum of the external region. The annular spectra were refitted using this new cosmic X-ray background model, and the change in the cluster temperature in each annulus relative to the standard fit was treated as a systematic uncertainty and added in quadrature to the statistical errors in each annulus.

To deproject and PSF-correct, we assume that the 3D temperature profile can be represented by a parametric model (adapted from Vikhlinin et al. 2006) that is convolved with a response matrix which simultaneously takes into account projection and PSF redistribution. This model was projected, taking into account the weighting scheme proposed by Vikhlinin (2006, see also Mazzotta et al. 2004) to correct for the bias introduced by fitting isothermal models to multi-temperature plasma emission, and fitted to the observed 2D annular temperature profile. Uncertainties were estimated from Monte Carlo randomisation of the projected temperature profile assuming a Gaussian distribution defined by the uncertainties on each data point, and then corrected to take into account the fact that parametric models tend to over-constrain the 3D profile. Full details of the depro-

<sup>2</sup>  $E(z)$  is the ratio of the Hubble constant at redshift  $z$  to its present value,  $E^2(z) = \Omega_M(1+z)^3 + \Omega_\Lambda$ .

<sup>3</sup> Our adoption of FWC data allows the use of a physical model for the X-ray background, in contrast to our previous analysis which used a blank sky background (Pratt et al. 2007). The results are consistent.



**Fig. 1.** Entropy profiles of the REXCESS sample, colour coded according to spectroscopic temperature measured in the  $[0.15-0.75] R_{500}$  aperture. Solid lines show the profiles derived from direct measurement; dotted lines show the entropy extrapolated into the central regions assuming an isothermal distribution at the temperature of the inner 3D data point (see text). On the left, the profiles are plotted in physical units; on the right, they are plotted in units of scaled radius  $R_{500}$  estimated from the  $M_{500}-Y_X$  relation given in Eqn. 1. In the right hand plot, dashed lines mark, from left to right, radii corresponding to  $0.1 R_{200}$ ,  $R_{2500}(\sim 0.45 R_{500})$  and  $R_{1000}(\sim 0.7 R_{500})$ .

jection and PSF correction of the temperature profiles, plus extensive tests of the robustness of the method, will be detailed in a forthcoming paper.

### 2.2.3. Entropy profiles

Since the density profiles are determined on a radial grid of significantly higher resolution than that of the temperature profiles, we determined the best fitting parametric 3D temperature profile on the same grid as that of the deprojected, deconvolved density profile and calculated the entropy,  $K = kT/n_e^{2/3}$ , accordingly.

In all cases, in the very central regions a single temperature bin encompasses a region covered by several density profile bins (the median number is 5). Given that the central density of the galaxy cluster population exhibits a dispersion of up to two orders of magnitude and the overall density profile changes by up to three orders of magnitude from the centre to the outskirts (e.g., Croston et al. 2008), while the temperature varies only by a factor of 2-3 (e.g., Pratt et al. 2007), it is clear that the characteristics of the density drive the properties of the entropy profiles. In order to examine the behaviour of the central entropy, we assume a constant central temperature, with the value given by the 3D temperature of the first shell. A similar procedure was used by Donahue et al. (2005) and Cavagnolo et al. (2009) and our adoption of this approach allows us to compare directly with their results. Note that for systems with poor central temperature profile resolution this extrapolation is only weakly model dependent, since it essentially concerns the disturbed systems, which have rather flat central temperature profiles (see Figure 3 of Arnaud et al. 2009). When this scheme is applied, the *XMM-Newton* profiles have a typical central resolution of

$\sim 5 h_{70}^{-1}$  kpc, which compares favourably with the typical resolution of  $\sim 2$  kpc in the *Chandra* analysis of Cavagnolo et al. (2009). The left hand panel of Figure 1 shows these entropy profiles plotted in physical units ( $h_{70}^{-1}$  kpc).

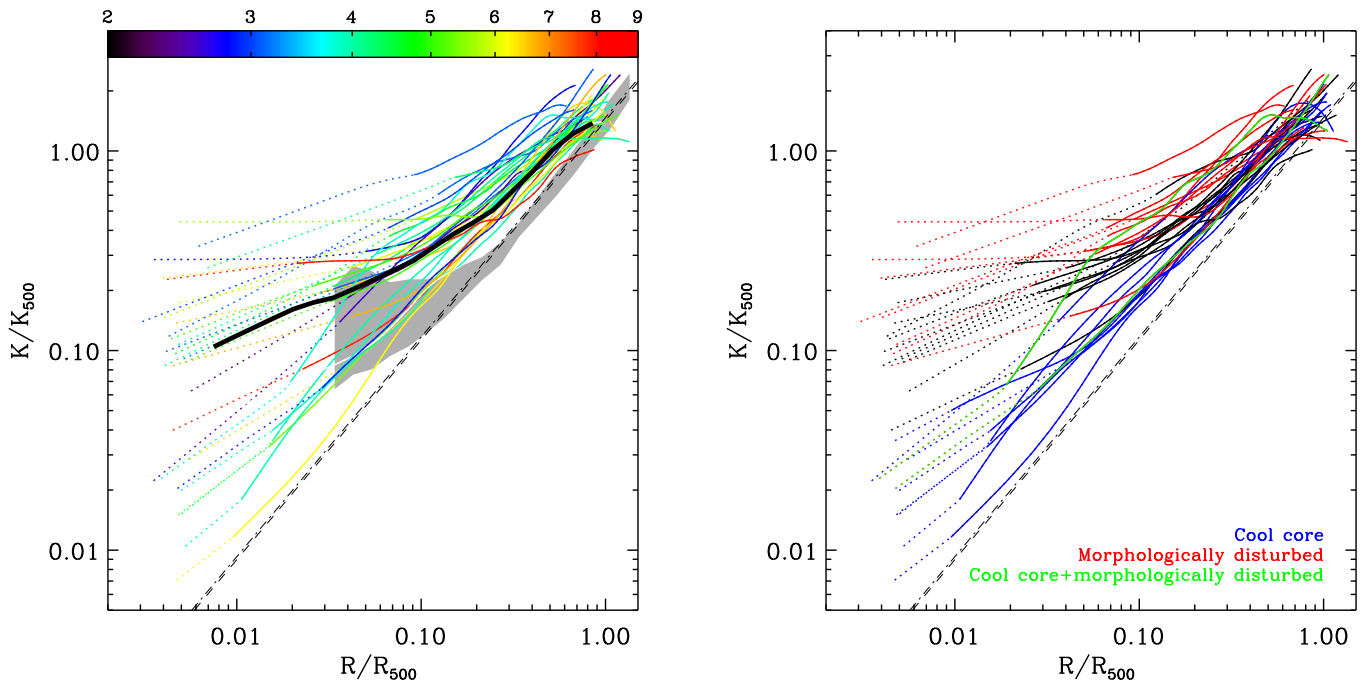
### 2.3. Scaling

In order to compare cluster profiles on a common radial scale, we express them in terms of  $R_\delta$ , the radius within which the mean mass density is  $\delta$  times the critical density at the cluster redshift<sup>4</sup>. For practical purposes, we generally scale to  $R_{500}$ , the effective limiting radius for high quality observations from *XMM-Newton* and *Chandra*. Since the sample contains systems in a variety of dynamical states, we use  $Y_X$  as a mass proxy. We estimate  $R_{500}$  iteratively as described in Kravtsov et al. (2006), from the updated calibration of the  $M_{500}-Y_X$  relation obtained by combining the Arnaud et al. (2007) results from nearby relaxed clusters with REXCESS data from morphologically relaxed systems. The full sample of 20 objects (8 from Arnaud et al. 2007 and 12 from REXCESS) is comprised of all systems for which the mass profiles are measured at least down to a density contrast  $\delta = 550$ . The resulting  $M_{500} - Y_X$  relation is:

$$E(z)^{2/5} M_{500} = 10^{14.567 \pm 0.010} \left[ \frac{Y_X}{2 \times 10^{14} M_\odot \text{ keV}} \right]^{0.561 \pm 0.018} h_{70}^{-1} M_\odot, \quad (1)$$

consistent with that derived by Arnaud et al. (2007) but with improved accuracy on slope and normalisation (Arnaud et al. 2009). We also use the spectroscopic temperature  $T$ , measured in the  $[0.15-0.75] R_{500}$  region, to investigate the scaling properties

<sup>4</sup>  $M_\delta = \delta \rho_c(z) (4\pi/3) R_\delta^3$ , where  $\rho_c(z) = E^2(z) 3H_0^2/8\pi G$ .



**Fig. 2.** Dimensionless entropy profiles of the REXCESS sample compared to theoretical expectations from non-radiative simulations. The observed profiles have been renormalised by the characteristic  $K_{500}$  as defined in Equation 3. Line styles are as for Figure 1. The dashed line depicts the best fitting power law fit to the the median entropy profile in the radial range  $[0.1 - 1]R_{200}$  for the clusters formed in the non-radiative simulations of Voit et al. (2005). The dot-dashed line represents the same relation corrected for a 13 per cent underestimate of true mass due to the use of hydrostatic equilibrium. *Left panel:* clusters arranged according to temperature. The thick black line represents the median of all observed dimensionless profiles. The shaded grey area corresponds to the region enclosed by the median profile and typical scatter of the SPH simulations in Voit et al. (2005). *Right panel:* subsamples identified, defined as in Section 2.1. Black profiles denote clusters that are neither cool core nor morphologically disturbed.

of the entropy and associated profiles. These values are given in Table 1.

The right hand panel of Figure 1 shows the entropy profiles plotted in terms of  $R_{500}$ . Plotting them this way explicitly shows the temperature dependence of the entropy distribution in the outer regions, and that despite the wide range of central entropy values, clusters clearly become more structurally similar with increasing radius. Beyond  $0.2 R_{500}$ , the relative dispersion in scaled entropy profiles is constant at approximately 30 per cent.

In the pure gravitational collapse scenario, the scaled profiles of any physical quantity should coincide, and so measures of these quantities at any scaled radius should correlate with global cluster parameters such as temperature or mass. The standard self-similar scalings are  $K \propto E(z)^{-4/3} T$  for temperature and  $K \propto E(z)^{-2/3} M^{2/3}$  for mass. Departures from these expectations are the direct result of the action of non-gravitational processes on the thermodynamics of the ICM. For comparison with previous work on entropy scaling relations, the entropy was measured for all clusters at radii equivalent to  $0.1 R_{200}$ ,  $R_{2500}$  ( $\approx 0.45 R_{500}$ ) and  $R_{1000}$  ( $\approx 0.72 R_{500}$ ), estimated from the scaling relations presented in Arnaud et al. (2005), via interpolation in the log-log plane. In addition, if  $R_{500}$  falls within the radial range encompassed by the centre of the outer temperature annulus (13 systems), we also calculated the entropy at  $R_{500}$ . Uncertainties on the entropy were estimated from the quadratic sum of the errors associated with the deconvolved density and temperature profiles. These values are listed in Table 1.

### 3. Comparison to theoretical expectations

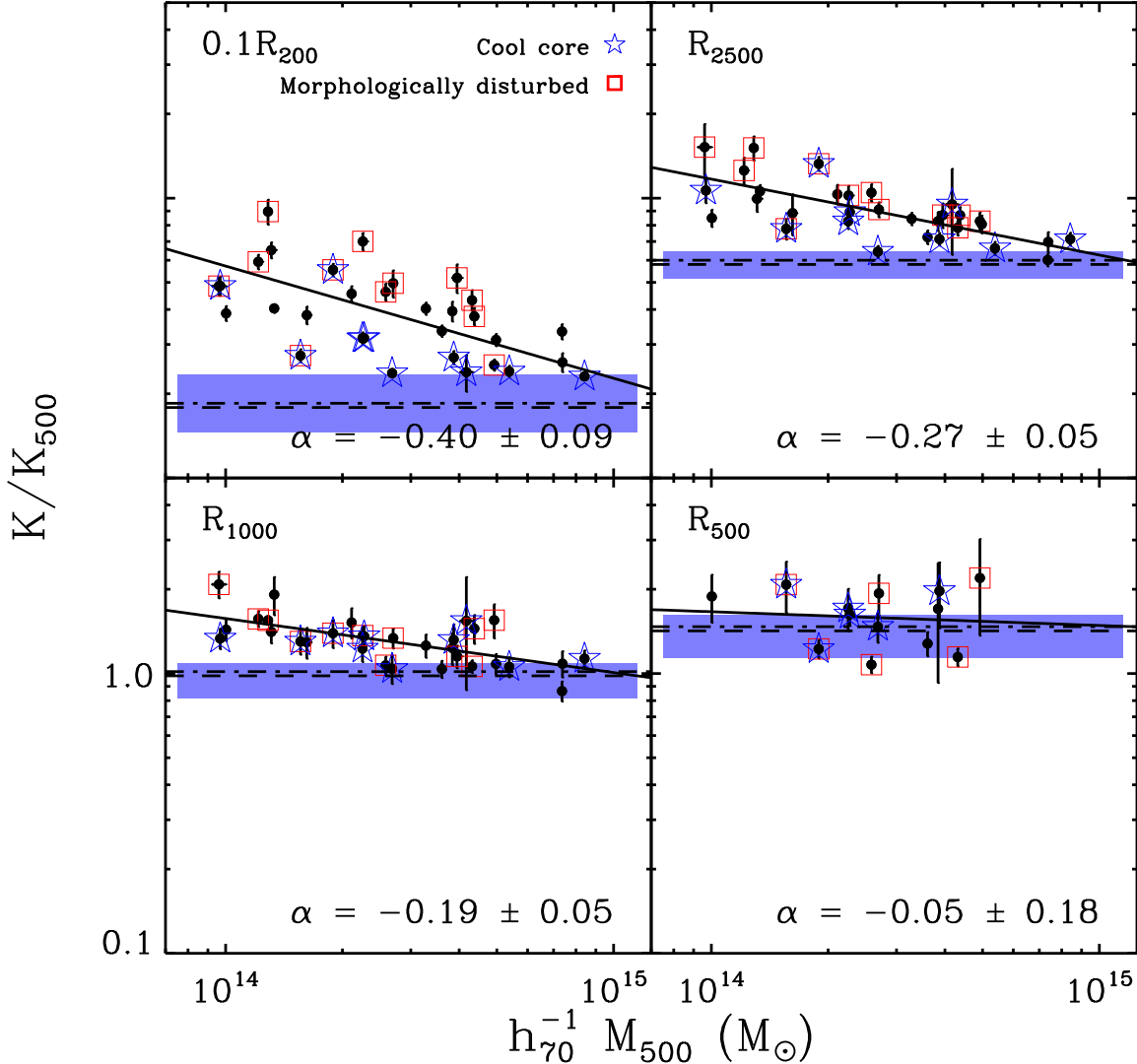
#### 3.1. Normalisation

We first address the question of the *absolute* entropy normalisation with respect to theoretical expectations in the case of purely gravitational structure formation. Any deviations from the predicted normalisation would point to the influence of non-gravitational processes on the thermodynamics of the ICM. Furthermore, radially resolved entropy profiles can be used to assess the radial extent of the change in ICM properties due to non-gravitational processes, if any. In the following, we compare the normalisation of our observed entropy profiles to theoretical expectations in order to elucidate the mass dependence and radial extent of these effects.

Numerical simulations which only implement gravitational processes make very specific predictions for both the normalisation and shape of galaxy cluster entropy profiles. Voit et al. (2005) discuss such simulations, pointing out that once scaled by the characteristic entropy of the halo,

$$K_{200} = \frac{1}{2} \left[ \frac{2\pi G^2 M_{200}}{15 f_b H(z)} \right]^{2/3}, \quad (2)$$

where  $f_b$  is the baryon fraction, the simulated SPH profiles, when fitted in the  $[0.1 - 1] R_{200}$  radial range, scatter about a median scaled profile described by the baseline power law relation  $K(R)/K_{200} = 1.32 (R/R_{200})^{1.1}$  with approximately 20 per cent dispersion. We note that non-radiative simulations do not predict power law profiles down to arbitrarily small radii; in fact,



**Fig. 3.** Dimensionless entropy  $K/K_{500}$  as a function of mass  $M_{500}$  for different fractions of scaled radius. The solid lines show the power law fit to the data at each scaled radius with exponent  $\alpha$ ; the dashed lines give the expected values from the power law relation  $K/K_{500} = 1.42 (R/R_{500})^{1.1}$  fitted to the non-radiative numerical simulations of Voit et al. (2005). The band indicates the typical dispersion of the simulated profiles about this relation at the radii indicated. The dot-dashed lines show the same relation adjusted for a 13 per cent mass underestimate due to the assumption of hydrostatic equilibrium. Cool core and morphologically disturbed subsamples are indicated. There is a clear segregation in central entropy properties of the subsamples at  $0.1 R_{200}$ , which disappears at larger scaled radius. The dimensionless entropy at  $R_{500}$  is slightly higher than, but approximately consistent with expectations from simple gravitational collapse, independent of mass.

significant flattening is generally found interior to  $R \lesssim 0.2 R_{500}$ . In addition, grid-based simulations consistently find larger entropy cores than SPH simulations even in the non-radiative case, which appears to be due to differences in particle mixing between the two different computational methods (Mitchell et al. 2009). However, beyond  $\sim 0.1 R_{200}$ , the agreement between simulations is better than 10 per cent (Voit et al. 2005; Mitchell et al. 2009). In the following, since we only compare at and beyond  $0.1 R_{200}$ , we compare to the SPH results only.

Assuming an abundance of  $Z = 0.3 Z_{\odot}$  and mean molecular weight  $\mu = 0.596$ , the characteristic entropy can be rewritten:

$$K_{500} = 106 \text{ keV cm}^{-2} \left( \frac{M_{500}}{10^{14} h_{70}^{-1} M_{\odot}} \right)^{2/3} \left( \frac{1}{f_b} \right)^{2/3} E(z)^{-2/3} h_{70}^{-4/3}. \quad (3)$$

For our assumed  $f_b = 0.15$  and  $R_{500}/R_{200} = 0.659$  derived from an NFW profile with  $c_{500} = 3.2$  (the mean  $c$  measured for a morphologically relaxed cluster sample by Pointecouteau et al. 2005), the baseline relation becomes

$$K(R)/K_{500} = 1.42 (R/R_{500})^{1.1}. \quad (4)$$

Note that this baseline relation was derived from simulations where the true masses were known, while our masses are calculated from an  $M-Y_X$  relation derived from hydrostatic mass estimates of morphologically relaxed systems. Piffaretti & Valdarnini (2008, and references therein) and Arnaud et al. (2009), applying hydrostatic mass estimates to a large number of simulated clusters, argue that such masses can be underestimated by  $-13 \pm 16$  per cent on average. In this case the normalisation factor in Eqn. 4 becomes 1.47.

**Table 2.** Best fitting parameters for the entropy-temperature and entropy-mass relations.

Radius	$C$ (keV cm <sup>-2</sup> )	$\alpha$	$\sigma_{\ln \kappa}$	
			raw	int
<u>Entropy-temperature relation</u>				
0.1 $R_{200}$	347 ± 23	0.89 ± 0.15	0.262 ± 0.040	0.254 ± 0.041
$R_{2500}$	783 ± 15	0.76 ± 0.06	0.120 ± 0.021	0.083 ± 0.118
$R_{1000}$	1152 ± 27	0.83 ± 0.06	0.093 ± 0.015	...
$R_{500}$	1489 ± 125	0.92 ± 0.24	0.265 ± 0.055	...
<u>Entropy-mass relation</u>				
$R_{2500}$	864 ± 27	0.42 ± 0.05	0.155 ± 0.025	0.136 ± 0.031
$R_{1000}$	1308 ± 52	0.48 ± 0.04	0.119 ± 0.017	0.052 ± 0.162
$R_{500}$	1748 ± 237	0.62 ± 0.17	0.265 ± 0.055	0.221 ± 0.160

$T$  is the spectroscopic temperature in the  $[0.15 - 0.75] R_{500}$  region; masses are estimated from the  $M - Y_X$  relation given in Eqn. 1. Data were fitted with a power law of the form  $E(z)^n K = C \times (A/A_0)^\alpha$ , with  $A_0 = 5$  keV and  $5.3 \times 10^{14} M_\odot$ , and  $n = 4/3$  and  $2/3$ , for temperature and mass respectively. Fits used orthogonal BCES regression with errors estimated using bootstrap resampling. The raw and intrinsic logarithmic scatter about the best fitting relations are given in the final two columns.

Figure 2 shows the dimensionless entropy profiles  $K(R)/K_{500}$ . Clearly, the observed profiles do not coincide. The central regions show the most dispersion and the profiles tend to converge towards the non-radiative prediction at large radius, but those of the lowest mass systems converge slowest; in other words, their slopes are shallower (discussed in more detail below). This is a manifestation of the fact that the entropy modification extends to larger radii in lower temperature systems, consistent with the expectation that non-gravitational processes have a greater effect at the low mass end of the cluster population.

To better quantify the above, Figure 3 shows the dimensionless entropy  $K/K_{500}$  versus mass for various fractions of scaled radius. Also overplotted in each panel is the expectation from Equation 4. As expected, it can be seen that at  $0.1 R_{500}$ , the excess with respect to the theoretical prediction from gravitational collapse is strongly mass-dependent, with the least massive systems exhibiting the strongest deviation. The mass dependence becomes less pronounced as we proceed towards the outer regions of the ICM, such that at  $R_{500}$ , the mass dependence is entirely consistent with zero and the measured values scatter about the theoretical prediction. Indeed, at this radius, the very slight negative slope can be attributed to the single lowest mass data point, which drives the fit.

At small radii there is large amount of scatter ( $\sigma_{\ln K, \text{int}} = 0.26 \pm 0.04$  about the best fitting regression line). Dividing the data into subsamples elucidates the origin of this scatter: there is a clear segregation in subsamples, with cool core systems showing the least deviation from the baseline prediction while morphologically disturbed systems show the most deviation. The subsample segregation disappears as we push outward though, and there is no evidence for any segregation at or beyond  $R_{1000}$ . The full radial behaviour of this trend is explicitly illustrated in the right hand panel of Figure 2.

Interior to  $R_{500}$ , the observed entropy is always higher than the baseline prediction. However, at  $R_{500}$ , the median dimensionless entropy is  $K(R_{500})/K_{500} = 1.70 \pm 0.35$ , where the uncertainty comes from the standard deviation of the points. This is slightly higher than, but consistent with both the baseline prediction of  $K(R_{500})/K_{500} = 1.42$  (Eqn. 4) and the same prediction corrected for a 13 per cent mass bias due to the assumption of hydrostatic

equilibrium,  $K(R_{500})/K_{500} = 1.47$ . The lack of mass dependence and agreement with the normalisation from simulations was also noted for a sample of cool core clusters by Nagai et al. (2007). For the present representative sample, given the large uncertainties it is not possible to test the predictions more thoroughly, underlining the need for robust, high quality, spatially resolved entropy measurements at and beyond  $R_{500}$ .

### 3.2. Entropy scaling relations

For comparison with previous work, we also examined the entropy-temperature and entropy-mass relations. Scaling relations were fitted with a power law of the form  $E(z)^n B = C(A/A_0)^\alpha$ , with  $A_0 = 5$  keV and  $5.3 \times 10^{14} M_\odot$  for  $T$  and  $M$  respectively, and  $n$  fixed to the expected scaling with redshift ( $n = 4/3$  for  $T$  and  $2/3$  for  $M$ ). Data were fitted using the orthogonal BCES minimisation technique (Akritas & Bershady 1996) with uncertainties on each fit parameter estimated from bootstrap resampling.

The best fitting slopes and intercepts for the entropy-temperature and entropy-mass relations at various scaled radii are listed in Table 2. The evolution of these slopes with increasing radius mirrors the behaviour of the dimensionless entropy discussed above; in the inner regions the relations are shallower than self-similar with large scatter, while at  $R_{500}$  the relations are compatible with self-similar (although with relatively large uncertainties given the limited number of data points).

Comparing to previous work, a wide variety of slopes have been found from fits to the entropy-temperature relation at  $0.1 R_{200}$ , ranging from very shallow ( $\alpha = 0.49 \pm 0.15$ : Pratt et al. 2006;  $\alpha = 0.50 \pm 0.08$ : Nagai et al. 2007) to very steep ( $\alpha = 0.92 \pm 0.12$ : Sanderson et al. 2009;  $\alpha = 0.85 \pm 0.19$ : this work). We simply note that cool core-only samples tend to yield shallower slopes than statistically-selected samples, a fact borne out in the present data, for which the entropy temperature relation at  $0.1 R_{200}$  for the cool core subsample has a slope of  $\alpha = 0.63 \pm 0.94$ , while the morphologically disturbed subsample has a slope of  $\alpha = 1.22 \pm 0.76$ . Beyond  $0.1 R_{200}$ , both the slope and the normalisation of the relations are in very good agreement with recent determinations (Nagai et al. 2007; Sun et al. 2009),

showing the excellent consistency between *XMM-Newton* and *Chandra* results<sup>5</sup>.

None of the studies listed above give constraints on entropy evolution. In this context, our results underline the need for representative samples, to establish the effects of non-gravitational processes and dynamical state on the evolution in the central regions. Furthermore, precise measurements at large radius ( $R \gtrsim R_{2500}$ ) are needed to establish the baseline entropy evolution in the absence of non-gravitational effects.

#### 4. Radial entropy structure

Various semi-analytical models and cosmological simulations of clusters formed in the absence of non-gravitational processes have shown that outside the central regions ( $R > 0.1 R_{200}$ , or  $\sim 0.15 R_{500}$ ), entropy profiles follow a power law with  $K(R) \propto R^{1.1}$  (Tozzi & Norman 2001; Borgani et al. 2005; Voit et al. 2002, 2005; Mitchell et al. 2009). Simulated profiles flatten in the very central regions due to entropy mixing (Wadsley et al. 2008; Mitchell et al. 2009). Observed profiles are also found generally to have similar external slopes (e.g., Pratt et al. 2006; Sun et al. 2009) and to flatten in the central regions in high resolution *Chandra* observations (Donahue et al. 2006; Sanderson et al. 2009; Cavagnolo et al. 2009). In the following, we investigate the central entropy slope and a parameterisation of the entropy profiles, and relate this to global cluster properties.

##### 4.1. Central slope

The left hand panel of Figure 4 shows the probability density distribution of the logarithmic slope  $d \ln K(R)/d \ln R$  of the entropy profiles measured at  $0.075 R_{500}$ , where we have directly measured data for all but two systems. The distribution shows two peaks, with the larger peak at a slope of  $\sim 0.4$ , containing approximately 2/3 of the sample, and a smaller peak at a slope of  $\sim 0.9$ , which comprises the cool core systems defined in the subsample classification scheme discussed in Section 2.1. Two peaks in the distribution of inner logarithmic slope were also found in the *Chandra* analysis of Sanderson et al. (2009, in their case, at  $0.05 R_{500}$ ).

##### 4.2. Parameterised fitting

###### 4.2.1. Model

A power law is the simplest parameterisation of the entropy profiles, but Figure 1 shows that such a description is clearly inadequate to describe the majority of the profiles in the present sample if the core regions are included in the fit. A possibly more interesting parameterisation is to instead model the entropy profiles with a power law plus a constant,  $K(R) = K_0 + K_{100} (R/100 h_{70}^{-1} \text{ kpc})^\alpha$  (as introduced by Donahue et al. 2005), where  $K_0$  represents the typical excess of core entropy above the best fitting power law at large radii. In this case, to ensure that the same radial range is fitted in all cases, and to ensure the precision of the powerlaw fit at large radii, the profiles were fitted between  $[R_{\text{in}} - R_{1000}]$ . The results of fitting the power law plus constant model are listed in Table 3. The external power law is

**Table 3.** Entropy profile parameterisation with a model of the form  $K(R) = K_0 + K_{100} (R/100 h_{70}^{-1} \text{ kpc})^\alpha$ . Entropy is measured in units of  $\text{keV cm}^{-2}$ .

Cluster	$K_0$	$K_{100}$	$\alpha$
RXC J0003.8+0203	39.33 ± 2.20	196.25 ± 3.93	0.77 ± 0.02
RXC J0006.6-3443	97.68 ± 26.02	280.10 ± 30.12	0.58 ± 0.05
RXC J0020.7-2542	177.45 ± 10.39	87.07 ± 9.76	1.15 ± 0.07
RXC J0049.4-2931	35.95 ± 2.32	141.70 ± 3.46	0.76 ± 0.02
RXC J0145.0-5300	240.05 ± 5.49	63.39 ± 5.40	1.40 ± 0.05
RXC J0211.4-4017	19.38 ± 1.58	118.87 ± 2.57	0.83 ± 0.02
RXC J0225.1-2928	94.57 ± 1.93	79.07 ± 3.74	1.54 ± 0.06
RXC J0345.7-4112	1.03 ± 0.44	153.59 ± 1.68	0.84 ± 0.01
RXC J0547.6-3152	141.30 ± 3.21	111.99 ± 4.49	1.08 ± 0.03
RXC J0605.8-3518	4.78 ± 0.67	139.62 ± 1.65	0.99 ± 0.01
RXC J0616.8-4748	22.46 ± 6.00	268.72 ± 9.09	0.58 ± 0.02
RXC J0645.4-5413	49.35 ± 2.65	145.40 ± 4.45	1.03 ± 0.02
RXC J0821.8+0112	...	263.95 ± 2.66	0.48 ± 0.01
RXC J0958.3-1103	24.78 ± 1.89	117.86 ± 3.58	1.15 ± 0.04
RXC J1044.5-0704	...	117.65 ± 0.53	0.95 ± 0.00
RXC J1141.4-1216	2.06 ± 0.17	148.75 ± 0.76	0.99 ± 0.01
RXC J1236.7-3354	52.14 ± 1.32	112.12 ± 2.18	1.04 ± 0.02
RXC J1302.8-0230	3.60 ± 0.42	209.91 ± 1.92	0.96 ± 0.01
RXC J1311.4-0120	49.20 ± 1.45	134.89 ± 2.10	1.10 ± 0.01
RXC J1516+0005	88.33 ± 3.68	165.48 ± 5.36	0.81 ± 0.02
RXC J1516.5-0056	43.01 ± 5.19	229.09 ± 7.74	0.66 ± 0.02
RXC J2014.8-2430	1.75 ± 0.26	121.61 ± 1.00	1.14 ± 0.01
RXC J2023.0-2056	52.27 ± 3.02	194.29 ± 4.41	0.69 ± 0.02
RXC J2048.1-1750	370.96 ± 6.31	14.53 ± 2.44	1.87 ± 0.09
RXC J2129.8-5048	150.49 ± 56.32	232.96 ± 56.20	0.54 ± 0.09
RXC J2149.1-3041	4.26 ± 0.69	136.22 ± 1.53	0.99 ± 0.01
RXC J2157.4-0747	92.24 ± 15.42	294.40 ± 19.04	0.52 ± 0.04
RXC J2217.7-3543	63.35 ± 3.25	155.35 ± 4.88	0.82 ± 0.02
RXC J2218.6-3853	98.10 ± 1.92	80.84 ± 2.91	1.36 ± 0.03
RXC J2234.5-3744	308.07 ± 6.98	62.01 ± 5.92	1.19 ± 0.05
RXC J2319.6-7313	6.18 ± 0.26	105.24 ± 0.93	0.98 ± 0.01

extremely stable to fitting ranges and clearly does not depend on the entropy of the central region or the value of  $K_{100}$ .

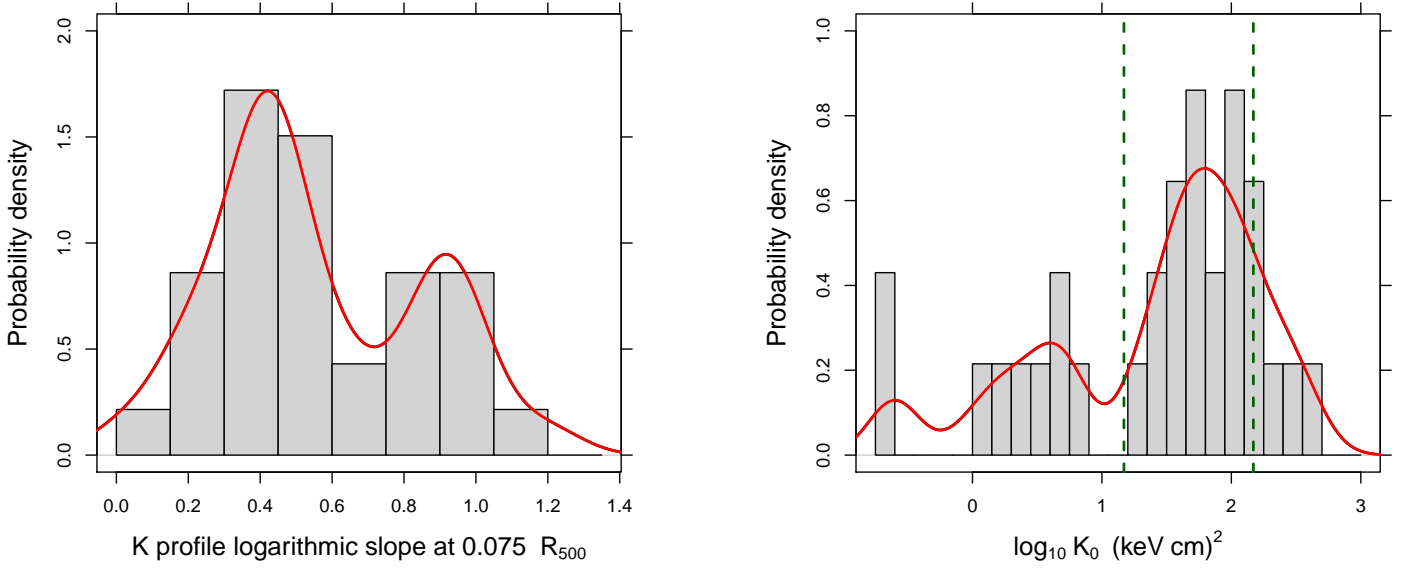
###### 4.2.2. Distribution of central entropy, $K_0$

The histogram of  $K_0$  values for the REXCESS sample is shown as a probability density in the right hand panel of Figure 4, where the bins are 0.15 dex and clusters with a  $K_0$  consistent with zero (i.e., pure power law profiles) are shown at the extreme left of the plot. For the REXCESS sample, the number of systems with a  $K_0$  consistent with zero at  $3\sigma$  significance is three for a constant central temperature (10 per cent of the sample), consistent with Cavagnolo et al. (2009), who analysed a large number of *Chandra* archive observations and who also used a constant central temperature assumption.

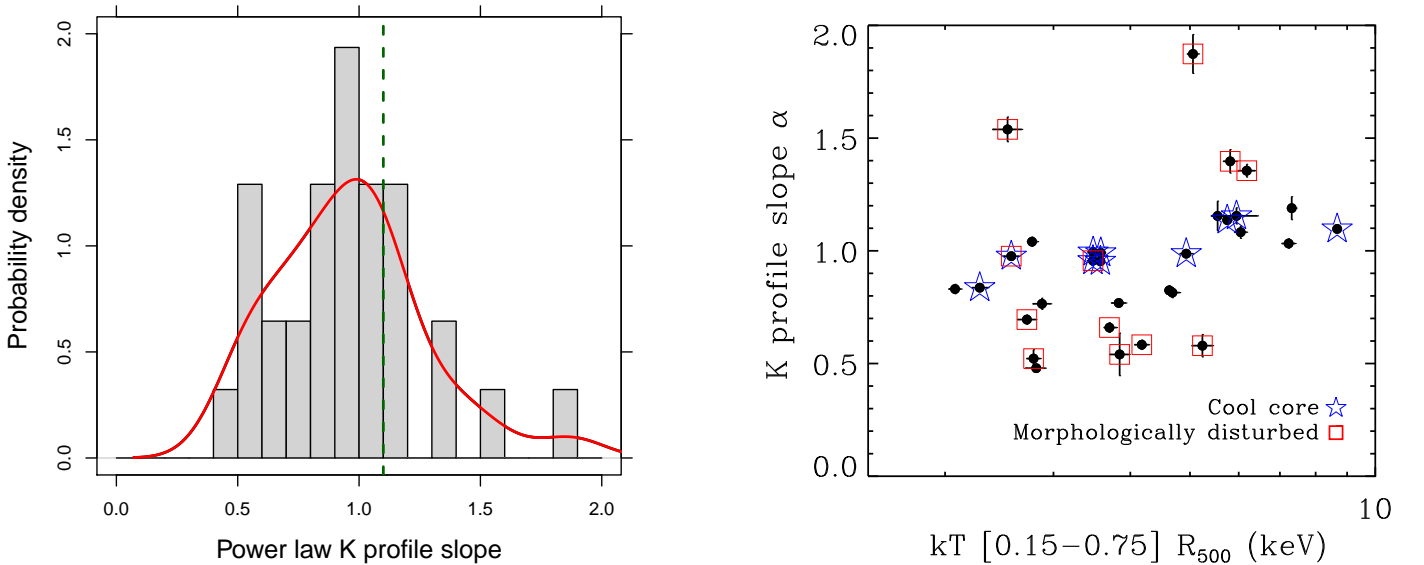
Cavagnolo et al. (2009) found that the distribution of central entropies in their sample is bimodal with peaks of approximately the same amplitude at  $K_0 \sim 15 \text{ keV cm}^2$  and  $K_0 \sim 150 \text{ keV cm}^2$ , and a distinct gap between  $K_0 \sim 30 - 50 \text{ keV cm}^2$  (see also Reiprich & Hudson 2007). The peak values are indicated in Figure 4 by a dashed line. The distribution of REXCESS central entropies also exhibits two peaks, although there are subtle differences in their positions and amplitudes. In particular, there are more clusters with a high  $K_0$  than with a low  $K_0$ , and the positions of both peaks are shifted somewhat to lower  $K_0$  with respect to those found by Cavagnolo et al.

<sup>5</sup> Our earlier results suggested somewhat shallower relations at  $R_{1000}$  (Pratt et al. 2006). The difference can be traced to the increased precision on the temperature profiles afforded by the present data, especially at low mass.





**Fig. 4.** Central regions of the REXCESS entropy profiles. *Left panel:* Probability density plot of the logarithmic slope of the density profile measured at a fiducial radius of  $0.075 R_{500}$ . Clusters with a steep slope correspond to cool core systems. The solid line is a kernel density plot with a smoothing width of 0.15. *Right panel:* Probability density plot of the central entropy excess above a power law,  $K_0$ , from a power law plus a constant model fit  $K(R) = K_0 + K_{100} (R/100 h_{70}^{-1} \text{ kpc})^\alpha$ , to the entropy profiles. The solid line is a kernel density plot with a smoothing width of 0.1. Cool core clusters have the smallest values of  $K_0$ .



**Fig. 5.** Outer regions of the REXCESS entropy profiles. *Left panel:* Probability density plot of outer slope values,  $\alpha$ , from a power law plus a constant model fit  $K(R) = K_0 + K_{100} (R/100 h_{70}^{-1} \text{ kpc})^\alpha$ , to the entropy profiles. The dashed line shows a value of 1.1. The solid line is a kernel density plot with a smoothing width of 0.1. *Right panel:* Entropy profile slope vs temperature in the  $[0.15 - 0.75] R_{500}$  aperture. Cool core and morphologically disturbed subsamples are indicated.

#### 4.2.3. Entropy slope outside the core

The left hand panel of Figure 5 shows the histogram of fitted values of the outer profile slope  $\alpha$  obtained from the power law plus constant model. While there is a quite substantial spread of values in entropy profile slope, ranging from extremely shallow ( $\alpha \sim 0.5$ ) to extremely steep ( $\alpha \sim 1.9$ ), there is no indication for bimodality in the distribution of outer entropy slope when the profiles are modelled in this way. The median slope for the power law plus constant model fits is 0.98, which is slightly shallower than the canonical value of 1.1 (indicated by the dashed line in the Figure). It is lower still than the value of 1.2 which was found by Voit et al. (2005) when fitting the  $[0.2 - 1] R_{200}$  radial range.

The right hand panel of Figure 5 shows the distribution of slopes versus temperature. The Spearman rank correlation coefficient is 0.53 with a significance of  $2.2 \times 10^{-3}$  indicating a significant correlation of slope with temperature. This is most likely a manifestation of the well-known dependence of outer density profile slope on temperature (see Croston et al. 2008, for the present sample). The different subsamples are indicated in the Figure by blue stars (cool core systems) and red squares (morphologically disturbed systems). Clearly the cool core subsample has a very small scatter in outer entropy slope; in addition there is a more pronounced trend with temperature, illustrated by the fact that the Spearman rank coefficient for this subsample is

0.77 with a significance of  $9.7 \times 10^{-3}$ . The morphologically disturbed subsample has a very large scatter, incorporating both the upper and lower extremes in outer slope values. The tight range of slopes in the cool core systems, together with the wide range of slopes in the morphologically disturbed systems, suggests that the core and external properties of the ICM are linked.

## 5. Discussion

### 5.1. Bimodality?

Two peaks are seen in the central regions of the REXCESS sample, which are visible both in the distribution of logarithmic entropy slope at  $0.075R_{500}$  and in the distribution of  $K_0$  from parameterisation of the profiles with a power law plus constant model (Figure 4). For a more quantitative comparison with the results of Cavagnolo et al. (2009), we performed a maximum likelihood fit of the 29 clusters for which  $K_0$  is constrained. The fitting was performed on the unbinned data in log space, using the MCLUST and FITDISTR packages in version 2.9 of the R statistical software environment<sup>6</sup>. We considered three different models: a single Gaussian, a left-skewed distribution, and a double Gaussian (i.e., a bimodal distribution), and used the Bayesian information criterion (BIC, Schwarz 1978) to distinguish between them. A difference in BIC of between 2 and 6 indicates positive evidence against the model with the greater BIC value; values above 6 indicate strong evidence against the model with the greater BIC value. The BIC values are -67.53, -63.32 and -63.08 for the single, left-skewed and double Gaussian distributions, respectively. Thus while the single Gaussian distribution is clearly the worst description of the data, on the basis of this test, and given the limited number of data points at our disposal, we cannot definitively distinguish between a bimodal and a left-skewed distribution of  $K_0$ .

The best bimodal distribution fit yields Gaussian means at  $K_0 \sim 3$  and  $K_0 \sim 75 \text{ keV cm}^2$ , with an amplitude ratio of 1:3. If we associate these values with the peaks found by Cavagnolo et al. (2009), then they are somewhat offset to lower values and the amplitude ratio is different (Cavagnolo et al. found peaks at  $\sim 15$  and  $\sim 150 \text{ keV cm}^2$ , with an amplitude ratio close to 1:1). The shift of the higher central entropy peak is due to a number of clusters in the REXCESS sample that fall directly in the  $K_0 \sim 30 - 50 \text{ keV cm}^2$  gap found by Cavagnolo et al. It may be that these clusters have a more ‘typical’ value of  $K_0$  for the general population. Given that the Cavagnolo et al. (2009) sample was archive limited, and given that prevailing sociological trends in cluster research have for many years led to a focus on extreme cool cores and spectacular mergers to the exclusion of more mundane objects, it may well be that clusters with a more ‘typical’ central entropy do not exist there but do exist in a representative sample. We note however that Cavagnolo et al. also detected bimodality in a complete flux-limited subsample of their data (although at much reduced significance), with the peaks shifted slightly lower relative to those of the full archive sample. Alternatively, the recent 2009 January *Chandra* calibration update may offer a more prosaic, albeit partial, explanation. Since the Cavagnolo et al. analysis predates this update, hotter systems with flat central temperature distributions ( $kT \gtrsim 4 \text{ keV}$ ), would have a systematically higher temperature, and hence entropy.

The shift of the lower central entropy peak appears to be due to technical differences connected to the treatment of temperature profiles. Cavagnolo et al. derived their entropy profiles from

*Chandra* data, which has sub-arcsecond resolution in the centre of the detector, meaning that no PSF correction was needed. However, while their density profiles were deprojected, their temperature profiles were not. For flat central temperature distributions, this will not substantially change the resulting entropy. However, for steeply declining central temperature distributions, neglect of projection effects will tend to lead to an overestimate of the temperature of the inner annulus (see e.g., Fig. 8 of Pointecouteau et al. 2004), and thus of the central entropy for the assumption of constant central temperature. The net effect of neglecting deprojection would be to shift the peak in  $K_0$  to higher values in the *Chandra* analysis, as observed.

We note that the differences in the central entropy distributions of REXCESS and Cavagnolo et al. (2009) are unlikely to be due to resolution effects. We extracted surface brightness profiles in  $3''.3$  bins while Cavagnolo et al. (2009) used  $5''$  bins, meaning that for a given redshift, our *XMM-Newton* entropy profiles in fact have a higher resolution than those from *Chandra*. Our  $3''.3$  bins give us a resolution of  $3.5\text{--}10 h_{70}^{-1} \text{ kpc}$  for the present sample, depending on redshift, with a median value of  $7 h_{70}^{-1} \text{ kpc}$ . Given that entropy cores extend typically to  $30 \text{ kpc}$  in *Chandra* observations of cool core systems (Donahue et al. 2005), these should be easily detectable with the current data if they exist. However, *Chandra* follow-up observations would still be desirable to quantify the cores of these systems at higher resolution, and particularly to investigate if there is indeed a turnover in entropy at very small radius in systems where we are unable to detect this effect.

We also do not find any bimodality in the distribution of outer power law slope  $\alpha$ , consistent with the increased self-similarity of the profiles in the outer regions. In this context it is important to underline the fact that our fitted model consisted of a power law plus an additive constant, and is thus different from the simple power law model fitted by e.g., Sanderson et al. (2009).

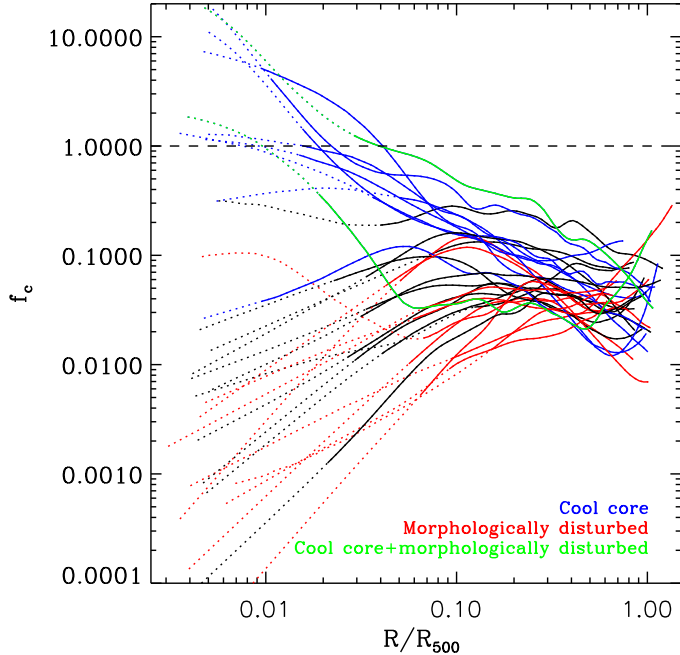
### 5.2. Maintaining the distribution of central entropies

The clear link between the lack of a cool core and the presence of morphological disturbance established by the REXCESS sample gives important clues to the processes at play in giving rise to the observed distribution of central entropies. It seems paradoxical that the profiles of cool core systems, where non-gravitational processes play an important role, resemble most closely the non-radiative baseline, while those of the unrelaxed objects, whose properties are expected to be dominated by gravitational processes, deviate the most from the baseline.

A possible explanation for the form of the cool core profiles is that in these systems, AGN heating is gentle and serves primarily to balance cooling, thus preserving the increasing form of the entropy profiles. If this were indeed the case, it would imply that AGN heating is achieved via mechanisms such as weak shocks (e.g., Fabian et al. 2003), or buoyant bubbles (e.g., Churazov et al. 2001), rather than via catastrophic explosions.

We can envisage two scenarios that could explain the properties of the morphologically disturbed clusters in REXCESS and the link between the lack of a cool core and the evidence that a cluster is unrelaxed. One explanation is that a combination of extra heating and dynamical activity have conspired to keep the central entropy elevated and so prevent the morphologically disturbed systems in the REXCESS from ever forming a cool core in the first place. This is possible because post merger disturbance persists for longer in these clusters, as in a disrupted core

<sup>6</sup> <http://www.r-project.org>



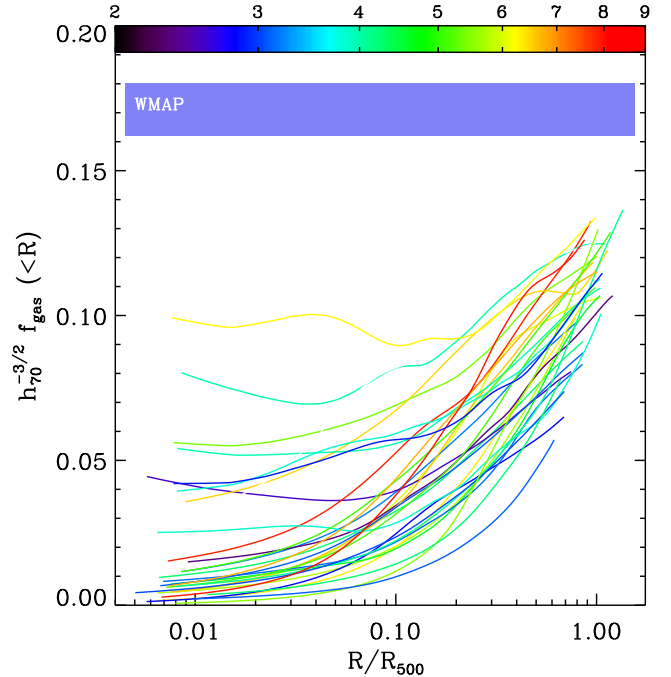
**Fig. 6.** Implied Spitzer conductivity suppression factor as a function of radius. The different cluster subsamples are indicated; line styles are as for Figure 1. The dashed line indicates the threshold between thermally unstable and conductively stable regimes.

with less entropy contrast, buoyancy differences will not be as strong, and the stratification that restores a cluster to a relaxed state will happen more slowly.

The second explanation is simply that the morphologically disturbed systems in the REXCESS were originally cool core systems, and that mergers have disrupted the entropy structure in the central regions. Given time, they should relax back to a cool core state, but we are seeing them before the cool core can re-establish itself. We note that such a scenario is supported in terms of timescales by the observed relationship between the lack of cool cores, clear evidence of dynamical disturbance and the presence of radio halo emission (e.g. Govoni et al. 2004). The ongoing REXCESS radio follow-up will be extremely useful in determining if this possibility can explain the clear anti-correlation of cool cores and disturbed morphology.

Furthermore, it is possible that the timescale for cooling to (re-) establish itself is impaired by conduction. Figure 6 shows the conduction suppression factor  $f_c$  expressed in terms of the radius (in kpc) and the entropy (in  $\text{keV cm}^2$ ), viz.,  $f_c \approx 62.5 (R^2/K^3)$  (Sanderson et al. 2009)<sup>7</sup>. It is clear that non-cool-core clusters can be stabilised since conductivity would need to be suppressed at most by only a factor of  $f_c \sim 0.1$  to locally counteract cooling (Figure 6), while this will not be effective in the centres of most cool core systems, where the threshold for conductive stability ( $f_c = 1$ ) is exceeded (a similar result was found by Sanderson et al. 2009). Thus conduction could also contribute to sustaining the elevated central entropy in post-merger systems long after the initial disruption.

<sup>7</sup> Note that this equation only holds at the limit of conductive thermal balance, i.e. where the radius exactly matches the Field length at all radii.



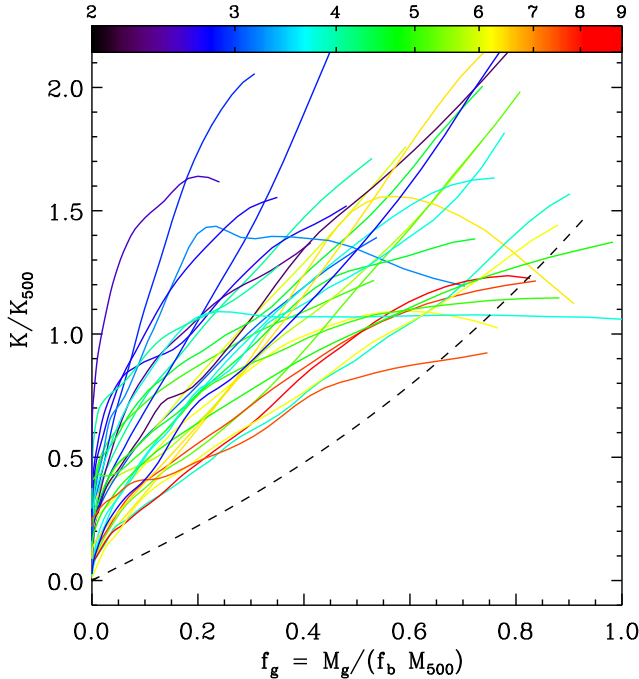
**Fig. 7.** Gas mass fraction profiles  $f_{\text{gas}}(<R)$ . Total mass profiles are calculated for an NFW profile with a concentration  $c_{500} = 3.2$  (Pointecouteau et al. 2005), normalised to the  $M_{500}$  estimated from Eqn. 1.

We note that high central entropy non-cool-core systems are difficult to reproduce in numerical simulations, which seems to be a consequence of the presence of very dense cores which are very hard to disrupt (e.g. Gómez et al. 2002; Poole et al. 2008). The presence of these dense cores, of which in some cases multiple instances may be present in the same system (Bóhringer et al., in preparation), at variance with observations, may point to deficiencies in the modelling of the complex interplay between gravitational and non-gravitational processes in these simulations.

### 5.3. Linking entropy and gas mass fraction

Entropy modification is generally discussed in terms of three basic mechanisms: early heating (‘pre-heating’), where the gas is heated before accretion into the dark matter potential well, presumably either by early supernovae and/or AGN activity (e.g. Kaiser 1991; Evrard & Henry 1991); internal heating after accretion by the same or similar mechanisms (e.g. Metzler & Evrard 1994; Bower et al. 2008); radiative cooling of the gas (e.g. Pearce et al. 2000), where the lowest entropy gas found in the centre of the cluster condenses and cools out of the ICM. All of these processes act to change the total amount of gas in the central regions of a cluster, either through making it more difficult to compress into the halo (early heating), through convection of gas to the outer regions (internal heating), or through physical removal of the gas to form stars (cooling).

In Figure 7 we show the gas mass fraction profiles  $f_{\text{gas}}(<R) = M_{\text{gas}}(<R)/M(<R)$  for the present sample. Gas masses have been calculated from the gas density profiles (Croston et al. 2008). Total mass profiles were calculated assuming an NFW profile with concentration  $c_{500} = 3.2$ , the average concentration

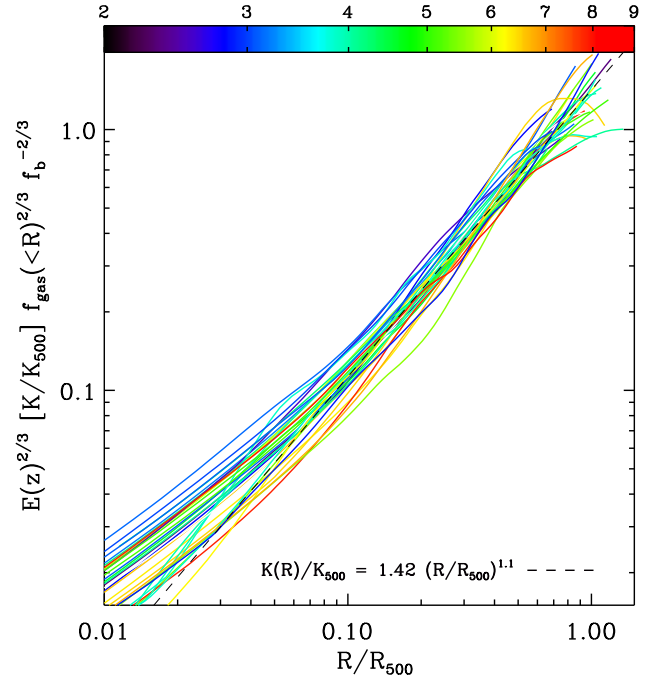


**Fig. 8.** The dimensionless entropy  $K/K_{500}$  rises with  $f_g(R) = M_g/(f_b M_{500})$ , the fraction of a cluster’s baryons in the ICM within radius  $R$  in all clusters. However, the entropy distribution shows a clear temperature dependence. The dashed line illustrates the unmodified entropy distribution for a cluster of  $M_{500} = 8 \times 10^{14} M_{\odot}$  (approximately the mass of the most massive cluster in the present sample), assuming a concentration  $c_{500} = 3.2$  and gas in hydrostatic equilibrium in the cluster potential with a density profile identical to that of the dark matter.

derived from the total mass profiles of the morphologically regular sample of Pointecouteau et al. (2005)<sup>8</sup>, normalised to  $M_{500}$  estimated from Eqn. 1. There is a clear dependence of  $f_{\text{gas}}$  on temperature/mass, throughout the observed temperature range, in the sense that hotter, more massive systems have higher gas mass fractions throughout the ICM. In addition, there is a clear dependence of gas mass fraction with radius in all systems, and only the most massive clusters have gas mass fractions which approach the universal value at the highest radii we are able to probe.

In Figure 8 we plot the dimensionless entropy  $K/K_{500}$  as a function of  $f_g(< R) = M_{\text{gas}}/(f_b M_{500})$ , the fraction of a cluster’s baryons in the ICM within radius  $R$ .  $f_g$  is calculated assuming  $f_b = 0.15$  ( $\Omega_b h^2 = 0.022$  and  $\Omega_m = 0.3$ ), and using total masses estimated from Eqn. 1. Overplotted for comparison is an *unmodified* entropy distribution from the models of Voit et al. (2002), derived for a cluster of  $M_{500} = 8 \times 10^{14} M_{\odot}$  (approximately the mass of the most massive cluster in the present sample), assuming a concentration  $c_{500} = 3$  and gas in hydrostatic equilibrium in the cluster potential with an identical density profile to that of the dark matter. This particular representation makes explicit both the dependence of the entropy distribution on baryon (gas) fraction, and the mass/temperature dependence of the baryon (gas) fraction itself. The gradual translation of the profiles to the left

<sup>8</sup> The dependence of concentration on total mass is negligible for the mass range we consider here (Pointecouteau et al. 2005; Buote et al. 2007).



**Fig. 9.** Entropy profiles multiplied by the gas mass fraction profile. The dashed line is the predicted entropy distribution from the non-radiative simulations of Voit et al. (2005).

hand side of the plot is due to a systematic lack of baryons (gas) in low temperature systems relative to high temperature systems.

The consequence of entropy modification is thus to remove gas (or prevent gas from accreting) in such a way as to leave both a radial *and* a mass dependence in the gas mass fraction. It is interesting to see whether correcting for this effect might bring the entropy profiles of our sample into agreement with the predictions from simulations. In Figure 9 we show the dimensionless entropy profiles multiplied by the gas mass fraction profile  $f_{\text{gas}}(< R)$ , a renormalisation that is equivalent to correcting simultaneously for the global *and* radial dependence of the gas mass fraction. Once renormalised in this way, the profiles are almost fully consistent, both in slope and normalisation, with the expectation from non-radiative simulations, and the dispersion drops dramatically. Slightly better agreement can be found if the simulated relation is multiplied by a factor to take into account the hydrostatic mass offset found in numerical simulations.

The above results can be explained in a number of ways. There may be a bias of gas accretion with mass, for instance due to early heating, which makes it more difficult to compress the gas into lower mass haloes. Once accreted, the gas may be removed from the hot phase by radiative cooling, which acts more efficiently in the densest central regions. However for this mechanism to be wholly responsible for the observed entropy properties would require it to affect the gas to a considerable fraction of  $R_{500}$  and to be preferentially efficient in low mass systems. Finally, there may be expulsion of material from the central regions towards the outskirts, perhaps via convection due to energy injection from supernovae or a central AGN, with the lowest mass systems experiencing the most central gas removal due to their shallower potential wells.

The mass dependence of the total baryon fraction provides a way of discriminating between the competing processes. If cool-

ing is the dominant effect, then the total baryon fraction should be almost constant across the mass range from groups to clusters as the low entropy gas is converted into stars. However, simply cooling out the low entropy gas would greatly exceed the observed mass in stars, and would lead to a galaxy luminosity function completely at odds with observations. If instead extra heating dominates, then the baryon fraction should be lower for lower mass, group-scale systems, as either early heating makes compression of gas into these haloes less efficient or AGN activity expels gas from their shallow potential wells. The observed anti-correlation in the relative dependencies of gas and stars with total mass implies that there is more mass locked in stars in systems which contain less gas (Lin et al. 2003; Gonzalez et al. 2007; Giodini et al. 2009). However, recent results suggest that there is still a  $\sim 3\sigma$  deficit of baryons with respect to that measured by WMAP on galaxy group scales (Giodini et al. 2009), implying that both cooling *and* heating must contribute to changing the thermodynamic properties of the ICM.

#### 5.4. Speculative scenario

The representative nature of the REXCESS sample has brought to light some intriguing points outlined above, which allow us to propose a tentative scenario. It seems that about two thirds of the REXCESS clusters possess a significantly higher central entropy than that expected from current non-radiative cosmological simulations and consequently do not possess a cooling core. A combination of extra heating and continuous ICM mixing due to merging may have kept these systems on a higher adiabat, leading to the observed high central entropies. Some early extra heating may have occurred in the protocluster phase, which would coincide with the peak in AGN activity at  $z \sim 2 - 3$ . In this scenario the lower entropy envelope traced by the non-cool core systems (see Figure 2) could indicate the level of early extra heating. The distribution of central entropy above this lower envelope would then be produced by later heating and gas mixing during mergers, with the least relaxed objects having the higher central entropy, as observed. These processes will inhibit formation of a stable cool core and naturally lead to redistribution of the gas to the outskirts, acting most importantly in low mass systems, corresponding to the observed behaviour of the gas mass fraction.

In contrast, about one third of the REXCESS sample possess a cool core. The clear association of the BCG with the bottom of the potential well (Haarsma et al. 2009) and their regular X-ray morphology testifies to the relaxed nature of these objects. These systems presumably experienced a less chaotic early dynamical history leading to a modest entropy elevation due to mixing (if any), and may have undergone less early extra heating, allowing them to develop a cool core at a relatively young age. The natural reduction of entropy due to cooling while the gas is still in the hot phase, due to the combination of a temperature drop and the consequent increase in gas density needed to keep pressure balance, may bring the profiles into line with the observed power law behavior. However, cooling must be limited by some finely-tuned feedback process to prevent a significant fraction of the gas from disappearing from the hot phase, which would lead to a net increase in entropy, at variance with the observations.

Our scenario bears some resemblance to the model proposed by McCarthy et al. (2008), although with increased emphasis on merger mixing as a process for setting and maintaining entropy levels in non-cool core systems.

## 6. Conclusions

Our data represent a considerable advance over those used in most previous analyses of the entropy structure and scaling in clusters. The sample of 31 clusters spans the temperature range [2-9] keV and includes systems with a variety of entropy characteristics. The objects have all been observed to approximately the same depth with the same instrument, allowing us to probe the properties of the entropy out to significant fractions of  $R_{500}$  (at least  $R_{1000}$  for all systems, and at least out to  $R_{500}$  for thirteen systems), which is essential to determine the radial extent of the effect of non-gravitational processes on the ICM.

In the inner regions, there is a mass dependent entropy excess with respect to theoretical expectations derived from cosmological numerical simulations including only gravity. At larger radii, the mass dependence weakens and the dispersion drops dramatically. The mass dependence disappears at  $R_{500}$ , and the entropy normalisation is, within the relatively large observational and theoretical uncertainties, in agreement with the expectations from non-radiative numerical simulations. This behaviour is mirrored in the entropy scaling relations, which are non-self similar at small radii but are compatible with self-similar at  $R_{500}$ . While similar results were found for a sample of cool core clusters by Nagai et al. (2007), it is important to note that in the group regime, at temperatures lower than 2 keV, the entropy normalisation of morphologically relaxed systems has been found to be significantly higher than predicted (Sun et al. 2009).

In the inner regions there is considerable dispersion, with a distinct segregation in residuals, in the sense that cool core clusters show the least deviation and morphologically disturbed systems show the most deviation with respect to expectations from non-radiative simulations. This dependence disappears at  $\sim R_{1000}$ . This clear association of unrelaxed morphology and elevated central entropy would suggest either that cool cores are destroyed by mergers, or that cool cores have never been able to form in these systems.

Fitting the entropy profiles with a power law plus constant model allows us to constrain  $\alpha$ , the power law slope at large radius, and also  $K_0$ , the central excess of entropy with respect to this power law. With the current data we cannot statistically distinguish between a bimodal distribution or a left-skewed distribution of  $K_0$  in log space; however, there is certainly no evidence for strong bimodality in the present sample. The distribution of outer entropy slopes is unimodal, with a median slope of 0.98. Cool cores have a narrow range of outer entropy slopes (0.8-1.2) while morphologically disturbed systems have a much wider range of outer slopes (0.5-1.9), suggesting a link between the properties of the cores and the outer regions of clusters.

In seeking to explain the structural and scaling behaviour of the entropy we looked at the gas mass fraction profiles of the sample. These are strongly mass dependent. Furthermore, the gas mass fraction increases with radius in all cases. A plot of the dimensionless entropy versus baryon fraction explicitly shows the dependence of entropy on gas mass. Renormalising the dimensionless entropy profiles by the gas mass fraction profile  $f_{\text{gas}}(< R)$ , effectively correcting simultaneously for both the mass and radial dependence of the gas mass fraction, dramatically decreases the dispersion in scaled profiles and brings them into agreement with predictions from non-radiative simulations. This provides further evidence for the underlying regularity of the cluster population, which has important implications for their use as cosmological probes.

The implication is that variations of gas content with mass and radius can explain the observed properties of the entropy

distributions, and by implication, the suppression of luminosity in low mass systems (e.g., Pratt et al. 2009). Various physical mechanisms can impose this behaviour, but the mass and radial dependence would strongly argue for a combination of both extra energy input and radiative cooling. However, it is necessary for these mechanisms to be capable of affecting the physical properties of the gas at least up to  $R_{1000}$ , and perhaps beyond, in order to explain the observed offsets with respect to expectations from non-radiative numerical simulations. We discuss a tentative scenario to explain the observed behaviour of the entropy and gas mass fraction in the REXCESS sample, based on a combination of extra heating and merger mixing maintaining elevated central entropy levels in the majority of the population, with a smaller fraction of systems able to develop a cool core.

In the near future, our ongoing REXCESS radio follow-up will help shed light on the relationship between merging activity and the lack of cool cores. Looking further ahead, observations of a similarly selected sample of galaxy groups, which are more sensitive to the effects of non-gravitational processes, would help to establish the magnitude of their impact at low masses. Precise, spatially-resolved measurements of the entropy at large radius ( $R > R_{500}$ ) will be a further essential test of theoretical models. Furthermore, a full census of the baryonic matter across the entire mass range for a *representative* sample of groups and clusters is essential to determine the relative contribution of the different processes in play.

*Acknowledgements.* We thank D. Nagai, A.J.R. Sanderson, M. Sun and A. Vikhlinin for helpful comments on the manuscript, and J. Ballet for useful discussions on statistical analysis. The present work is based on observations obtained with *XMM-Newton*, an ESA science mission with instruments and contributions directly funded by ESA Member States and the USA (NASA). GWP and HB acknowledge partial support from the DfG Transregio project TR33, "Dark Universe".

## References

- Akritas, M. G. & Bershad, M. A. 1996, *ApJ*, 470, 706  
 Arnaud, M., Pointecouteau, E., & Pratt, G. W. 2005, *A&A*, 441, 893  
 Arnaud, M., Pointecouteau, E., & Pratt, G. W. 2007, *A&A*, 474, L37  
 Arnaud, M., Pratt, G. W., Piffaretti, R., et al. 2009, *A&A*, submitted (arXiv:0910.1234)  
 Böhringer, H., Schuecker, P., Guzzo, L., et al. 2004, *A&A*, 425, 367  
 Böhringer, H., Schuecker, P., Pratt, G. W., et al. 2007, *A&A*, 469, 363  
 Borgani, S., Finoguenov, A., Kay, S. T., et al. 2005, *MNRAS*, 361, 233  
 Borgani, S. & Kravtsov, A. 2009, *Advanced Science Letters*, in press (arXiv:0906.4370)  
 Bower, R. G., McCarthy, I. G., & Benson, A. J. 2008, *MNRAS*, 390, 1399  
 Buote, D. A., Gastaldello, F., Humphrey, P. J., et al. 2007, *ApJ*, 664, 123  
 Cavagnolo, K. W., Donahue, M., Voit, G. M., & Sun, M. 2009, *ApJS*, 182, 12  
 Churazov, E., Brügggen, M., Kaiser, C. R., Böhringer, H., & Forman, W. 2001, *ApJ*, 554, 261  
 Croston, J. H., Arnaud, M., Pointecouteau, E., & Pratt, G. W. 2006, *A&A*, 459, 1007  
 Croston, J. H., Pratt, G. W., Böhringer, H., et al. 2008, *A&A*, 487, 431  
 David, L. P., Jones, C., & Forman, W. 1996, *ApJ*, 473, 692  
 De Luca, A. & Molendi, S. 2004, *A&A*, 419, 837  
 Donahue, M., Horner, D. J., Cavagnolo, K. W., & Voit, G. M. 2006, *ApJ*, 643, 730  
 Donahue, M., Voit, G. M., O'Dea, C. P., Baum, S. A., & Sparks, W. B. 2005, *ApJ*, 630, L13  
 Ettori, S., Tozzi, P., Borgani, S., & Rosati, P. 2004, *A&A*, 417, 13  
 Evrard, A. E. & Henry, J. P. 1991, *ApJ*, 383, 95  
 Fabian, A. C., Sanders, J. S., Allen, S. W., et al. 2003, *MNRAS*, 344, L43  
 Finoguenov, A., Jones, C., Böhringer, H., & Ponman, T. J. 2002, *ApJ*, 578, 74  
 Giodini, S., Pierini, D., Finoguenov, A., et al. 2009, *ApJ*, 703, 982  
 Gómez, P. L., Loken, C., Roettiger, K., & Burns, J. O. 2002, *ApJ*, 569, 122  
 Gonzalez, A. H., Zaritsky, D., & Zabludoff, A. I. 2007, *ApJ*, 666, 147  
 Govoni, F., Markevitch, M., Vikhlinin, A., et al. 2004, *ApJ*, 605, 695  
 Haarsma, D. B., Leisman, L., Donahue, M., et al. 2009, *ApJ*, submitted (arXiv:0911.2798)  
 Jeltema, T. E., Mulchaey, J. S., Lubin, L. M., Rosati, P., & Böhringer, H. 2006, *ApJ*, 649, 649  
 Johnson, R., Ponman, T. J., & Finoguenov, A. 2009, *MNRAS*, 395, 1287  
 Kaiser, N. 1991, *ApJ*, 383, 104  
 Kravtsov, A. V., Vikhlinin, A., & Nagai, D. 2006, *ApJ*, 650, 128  
 Lin, Y.-T., Mohr, J. J., & Stanford, S. A. 2003, *ApJ*, 591, 749  
 Lloyd-Davies, E. J., Ponman, T. J., & Cannon, D. B. 2000, *MNRAS*, 315, 689  
 Lumb, D. H., Warwick, R. S., Page, M., & De Luca, A. 2002, *A&A*, 389, 93  
 Mazzotta, P., Rasia, E., Moscardini, L., & Tormen, G. 2004, *MNRAS*, 354, 10  
 McCarthy, I. G., Babul, A., Bower, R. G., & Balogh, M. L. 2008, *MNRAS*, 386, 1309  
 Metzler, C. A. & Evrard, A. E. 1994, *ApJ*, 437, 564  
 Mitchell, N. L., McCarthy, I. G., Bower, R. G., Theuns, T., & Crain, R. A. 2009, *MNRAS*, 395, 180  
 Morandi, A. & Ettori, S. 2007, *MNRAS*, 380, 1521  
 Nagai, D., Kravtsov, A. V., & Vikhlinin, A. 2007, *ApJ*, 668, 1  
 Pearce, F. R., Thomas, P. A., Couchman, H. M. P., & Edge, A. C. 2000, *MNRAS*, 317, 1029  
 Piffaretti, R., Jetzer, P., Kaastra, J. S., & Tamura, T. 2005, *A&A*, 433, 101  
 Piffaretti, R. & Valdarnini, R. 2008, *A&A*, 491, 71  
 Pointecouteau, E., Arnaud, M., Kaastra, J., & de Plaa, J. 2004, *A&A*, 423, 33  
 Pointecouteau, E., Arnaud, M., & Pratt, G. W. 2005, *A&A*, 435, 1  
 Ponman, T. J., Cannon, D. B., & Navarro, J. F. 1999, *Nature*, 397, 135  
 Ponman, T. J., Sanderson, A. J. R., & Finoguenov, A. 2003, *MNRAS*, 343, 331  
 Poole, G. B., Babul, A., McCarthy, I. G., Sanderson, A. J. R., & Fardal, M. A. 2008, *MNRAS*, 391, 1163  
 Pratt, G. W. & Arnaud, M. 2003, *A&A*, 408, 1  
 Pratt, G. W., Arnaud, M., & Pointecouteau, E. 2006, *A&A*, 446, 429  
 Pratt, G. W., Böhringer, H., Croston, J. H., et al. 2007, *A&A*, 461, 71  
 Pratt, G. W., Croston, J. H., Arnaud, M., & Böhringer, H. 2009, *A&A*, 498, 361  
 Reiprich, T. H. & Hudson, D. S. 2007, in *Heating versus Cooling in Galaxies and Clusters of Galaxies*, ed. H. Böhringer, G. W. Pratt, A. Finoguenov, & P. Schuecker, 381–+  
 Sanderson, A. J. R., O'Sullivan, E., & Ponman, T. J. 2009, *MNRAS*, 395, 764  
 Schwarz, G. 1978, *Ann. Stat.*, 6, 461  
 Sun, M., Voit, G. M., Donahue, M., et al. 2009, *ApJ*, 693, 1142  
 Tozzi, P. & Norman, C. 2001, *ApJ*, 546, 63  
 Vikhlinin, A. 2006, *ApJ*, 640, 710  
 Vikhlinin, A., Kravtsov, A., Forman, W., et al. 2006, *ApJ*, 640, 691  
 Voit, G. M. 2005, *Reviews of Modern Physics*, 77, 207  
 Voit, G. M., Bryan, G. L., Balogh, M. L., & Bower, R. G. 2002, *ApJ*, 576, 601  
 Voit, G. M., Kay, S. T., & Bryan, G. L. 2005, *MNRAS*, 364, 909  
 Wadsley, J. W., Veeravalli, G., & Couchman, H. M. P. 2008, *MNRAS*, 387, 427  
 Zhang, Y.-Y., Finoguenov, A., Böhringer, H., et al. 2008, *A&A*, 482, 451