

HALO GAS CROSS SECTIONS AND COVERING FRACTIONS OF Mg II ABSORPTION SELECTED GALAXIES

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ABSTRACT

We examine halo gas cross sections and covering fractions, f_c , of intermediate-redshift Mg II absorption selected galaxies. We computed statistical absorber halo radii, R_x , using current values of dN/dz and Schechter luminosity function parameters, and have compared these values to the distribution of impact parameters and luminosities from a sample of 37 galaxies. For equivalent widths $W_r(2796) \geq 0.3 \text{ \AA}$, we find $43 \leq R_x \leq 88 \text{ kpc}$, depending on the lower luminosity cutoff and the slope, β , of the Holmberg-like luminosity scaling, $R \propto L^\beta$. The observed distribution of impact parameters, D , are such that several absorbing galaxies lie at $D > R_x$ and several non-absorbing galaxies lie at $D < R_x$. We deduced that f_c must be less than unity and obtain a mean of $\langle f_c \rangle \sim 0.5$ for our sample. Moreover, the data suggest that halo radii of Mg II absorbing galaxies do not follow a luminosity scaling with β in the range of 0.2–0.28, if $f_c = 1$ as previously reported. However, provided $f_c \sim 0.5$, we find that halo radii can remain consistent with a Holmberg-like luminosity relation with $\beta \simeq 0.2$ and $R_* = R_x/\sqrt{f_c} \sim 110 \text{ kpc}$. No luminosity scaling ($\beta = 0$) is also consistent with the observed distribution of impact parameters if $f_c \leq 0.37$. The data support a scenario in which gaseous halos are patchy and likely have non-symmetric geometric distributions about the galaxies. We suggest that halo gas distributions may not be governed primarily by galaxy mass/luminosity but also by stochastic processes local to the galaxy.

Key words: galaxies: halos – galaxies: ISM – quasars: absorption lines

1. INTRODUCTION

Understanding galaxy formation and evolution is one of the most important topics of modern astronomy. The extended distribution of baryonic gas surrounding galaxies holds great potential for constraining theories of their formation. However, the sizes of gaseous galaxy halos along with the distribution of gas within are not well understood. Numerical models have been able to synthesize the formation and evolution of large scale structures; however, there are unresolved issues regarding the evolution of individual galaxies and halos. The halo baryon-fraction problem (e.g., Mo & Mao 2002) and the rapid cooling of gas (e.g., White & Rees 1978) result in galaxy halos which have little or no gas soon after they form. These effects are not seen in the observable universe since there is an abundance of galaxies where gas has been detected in halos via quasar absorption lines.

From an observational standpoint, quasar absorption lines provide a unique means of probing the extent and abundance of halo gas. Although quasar absorption line observations to date are sufficient to recognize the aforementioned problems, they are lacking the details required to statistically constrain the distribution of the baryonic gas in the halos of simulated galaxies. Cross-correlations between absorbers and galaxies hold the promise to yield useful information on cloud sizes and halo gas covering fractions. First steps toward incorporating multi-phase gas in semi-analytical models and numerical simulations suggest that warm gas in halos extends out to galactocentric distances of $\sim 150 \text{ kpc}$ with cloud covering fractions of ~ 0.25 – 0.6 (Maller & Bullock 2004; Kaufmann et al. 2006).

The association of Mg II $\lambda\lambda 2796, 2803$ doublet absorption in quasar spectra with normal, bright, field galaxies has been firmly established (e.g., Bergeron & Boissé 1991; Steidel et al. 1994; Churchill et al. 2005). In an effort to understand halo sizes

and gas distributions, Steidel (1995, hereafter S95) searched for foreground galaxies associated with Mg II absorption within $\sim 10''$ ($\sim 65 \text{ kpc}$ for $z = 0.5$) of quasars.⁵ The sample consisted of 53 absorbing and 14 non-absorbing galaxies with a Mg II $\lambda 2796$ equivalent width sensitivity limit of $W_r(2796) > 0.3 \text{ \AA}$. S95 directly fitted the data by assuming a Holmberg-like luminosity scaling,

$$R(L) = R_* \left(\frac{L}{L_*} \right)^\beta \text{ kpc}, \quad (1)$$

and minimizing the number of non-absorbing and absorbing galaxies above and below the $R(L)$ relation. The best fit obtained clearly showed that absorbing and non-absorbing galaxies could be separated and that the halo radii $R(L_K)$ and $R(L_B)$ scale with luminosity with $\beta = 0.15$ and $\beta = 0.2$, respectively, where an L_B^* galaxy has a gas halo cross section of $R_* = 55 \text{ kpc}$. Furthermore, since almost none of the absorbing galaxies were observed above the $R(L)$ boundary and that almost none of the non-absorbing galaxies were observed below the $R(L)$ boundary, S95 inferred that *all* $L > 0.05L^*$ galaxies are hosts to Mg II absorbing gas halos characterized by a covering fraction of unity and a spherical geometry which truncates at $R(L)$. Examination of this now “standard model” has been the subject of several theoretical studies (e.g., Charlton & Churchill 1996; Mo & Miralda-Escude 1996; Lin & Zou 2001).

Guillemin & Bergeron (1997) determined a steeper value of $\beta = 0.28$ for the B -band luminosity obtained from a best fit to the upper envelope of the distribution of impact parameters of 26 absorbing galaxies. They found $R_* = 67 \text{ kpc}$.

⁵ Throughout we adopt a $h = 0.70$, $\Omega_M = 0.3$, $\Omega_\Lambda = 0.7$ cosmology. All quoted physical quantities from previously published works have been converted to this cosmology.

Using a reverse approach of establishing foreground galaxy redshifts and then searching for Mg II absorption in the spectra of background quasars yields results inconsistent with a covering fraction of unity. For example, Bowen et al. (1995) identified 17 low-redshift galaxies with background quasar probing an impact parameter range between 3 and 162 kpc. Galaxies that were probed at impact parameters greater than 13 kpc had no absorption in the halo ($W_r(2796) \geq 0.40\text{--}0.9 \text{ \AA}$); however, four of the six galaxies within 13 kpc of the halo produced Mg II absorption. For intermediate-redshift galaxies, Bechtold & Ellingson (1992) reported a covering fraction $f_c \simeq 0.25$ for $W_r(2796) \geq 0.26 \text{ \AA}$ for eight galaxies with $D \leq 85$ kpc. Also, Tripp & Bowen (2005) reported $f_c \sim 0.5$ for $W_r(2796) \geq 0.15 \text{ \AA}$ for ~ 20 galaxies with $D \leq 50$ kpc. These results are also consistent with the findings of Churchill et al. (2005) who reported very weak Mg II absorption, $W_r(2796) < 0.3 \text{ \AA}$, well inside the $R(L)$ boundary of bright galaxies; these galaxies would be classified as “non-absorbers” in previous surveys. They also report $W_r(2796) > 1 \text{ \AA}$ absorption out to $\simeq 2R(L)$. All these results suggest that there are departures from the standard model, that the covering fraction of Mg II-absorbing gas is less than unity, and that the halo sizes and the distribution of the gas appear to diverge from the $R(L)$ relation with spherical geometry.

Another approach to understanding halo sizes and gas distributions is to determine the statistical properties of Mg II absorbing gas and then compute the statistical cross section from the redshift path density, dN/dz (see Lanzetta et al. 1995). The downfall of this method is that a galaxy luminosity function must be adopted in order to estimate R_* . Nestor et al. (2005) acquired a sample of over 1300 Mg II absorption systems, with $W_r(2796) \geq 0.3 \text{ \AA}$ from the Sloan Digital Sky Survey (SDSS). Using the K -band Holmberg-like luminosity scaling and luminosity function of MUNICS (Drory et al. 2003), Nestor et al. computed $R_* = 60\text{--}100$ kpc for adopted minimum luminosity cutoffs of $L_{\min} = 0.001\text{--}0.25L^*$. They found no redshift evolution of R_* over the explored range of $0.3 \leq z \leq 1.2$.

Zibetti et al. (2007) studied the statistical photometric properties of ~ 2800 Mg II absorbers in quasar fields imaged with SDSS. Using the method of image stacking, they detected low-level surface brightness (SB) azimuthally about the quasar. The SB profiles follow a decreasing power law with projected distance away from the quasar out to 100–200 kpc. These results imply that absorption selected galaxies may reside out to projected distances of 200 kpc. However, it is worth noting that the extended light profiles may be an artifact of clustering of galaxies. Cluster companions of the Mg II absorbing galaxies could extend the observed light profile over hundreds of stacked images. Thus, one would infer that Mg II absorbing galaxies are present at larger impact parameters than would be found in direct observation of individual galaxies.

Motivated by recent expectations from simulations that halo gas is dynamically complex and sensitive to the physics of galaxy formation, we investigate the standard halo model of Mg II absorbers. We also aim to provide updated constraints on f_c and β for galaxy formation simulations. In this paper, we demonstrate that $f_c < 1$ and question the validity of the Holmberg-like luminosity scaling (Equation (1)). Using high resolution quasar spectra, we explore Mg II absorption strengths to an order of magnitude more sensitive than previous surveys which allow us to re-identify non-absorbing galaxies as “weak” absorbing galaxies. In Section 2, we describe our sample and analysis. In Section 3, we present new calculations

of the statistical absorber radius computed using the statistically measured absorption path density dN/dz and the Schechter luminosity function. We then compare these values to the empirical results of S95 and to a sample of known Mg II absorption selected galaxies with measured luminosities and impact parameters. We also examine how individual halos behave with respect to the statistical halo. In Section 4, we discuss the properties and distribution of gas in halos. Our concluding remarks are in Section 5.

2. DATA AND ANALYSIS

We have constructed a sample of 37 galaxies ($0.3 < z < 1.0$), with spectroscopically confirmed redshifts, selected by the presence of Mg II absorption in quasar spectra. The absorption properties were measured from HIRES/Keck (Vogt et al. 1994) and UVES/VLT (Dekker et al. 2000) spectra. The Mg II $\lambda 2796$ profiles have been presented in Churchill et al. (2005), where the detection limit is $W_r(2796) \geq 0.02 \text{ \AA}$ (5σ). Galaxy properties were measured from F702W or F814W WFPC-2/HST images of the quasar fields. Images of the galaxies, along with further details of the sample selection, data, and data analysis, can be found in Kacprzak et al. (2007a).

Galaxy absolute magnitudes, M_B , were determined from the k -corrected observed m_{F702W} or m_{F814W} adopted from Kacprzak et al. (2007a). The k -corrections were computed using the formalism of Kim et al. (1996) based upon the spectral energy distribution (SED) templates of Kinney et al. (1996). The adopted SED for each galaxy was based upon its rest-frame $B - K$ color (SDP94). For galaxies with no color information, we adopted a Sb SED which is consistent with average color of Mg II absorbing galaxies (SDP94; Zibetti et al. 2007). Our k -corrections are consistent with those from the literature (Kim et al. 1996; Fukugita et al. 1995). B -band luminosities were computed using the DEEP2 optimal M_B^* of Faber et al. (2007, Table 2) in the redshift bin appropriate for each galaxy. M_B^* ranges from -21.07 ($\langle z \rangle = 0.3$) to -21.54 ($\langle z \rangle = 1.1$).

We compute the halo gas cross section determined from the redshift path density,

$$\frac{dN}{dz} = \pi R_x^2 \cdot \Phi^* \Gamma(x, y) \cdot \frac{c}{H_0} \frac{(1+z)^2}{\sqrt{\Omega_m(1+z)^3 + \Omega_\Lambda}}, \quad (2)$$

where R_x is the statistical absorber radius for an L^* galaxy, and Φ^* is the number density of L^* galaxies. $R_x^2 = f_c R_*^2$, where R_* is the covering fraction corrected absorbing halo radius. Note here that we make a distinction between R_x , which is derived from the redshift path density, and R_* , which is a physical cross section of the absorbing gas accounting for the covering fraction. $\Gamma(x, y)$ is the incomplete Gamma function in which $x = 2\beta - \alpha + 1$, where α is the faint-end slope of the Schechter galaxy luminosity function and β parameterizes a Holmberg-like luminosity scaling of Equation (1). The parameter $y = L_{\min}/L^*$, where L_{\min} is the minimum luminosity of galaxies contributing to absorption. The influence of y on the value of R_x becomes relatively more important as $\beta \rightarrow 0$.

We present our sample in Figure 1(a), plotting $W_r(2796)$ versus L_B/L_B^* . The solid points have $W_r(2796) \geq 0.3 \text{ \AA}$ and the open points are weak systems (Churchill et al. 1999), having $W_r(2796) < 0.3 \text{ \AA}$, and would have been classified as non-absorbing galaxies in previous surveys (e.g., SDP94; Guillemin & Bergeron 1997). Since R_x is computed using dN/dz which is determined for systems with $W_r(2796) \geq 0.3 \text{ \AA}$, we must

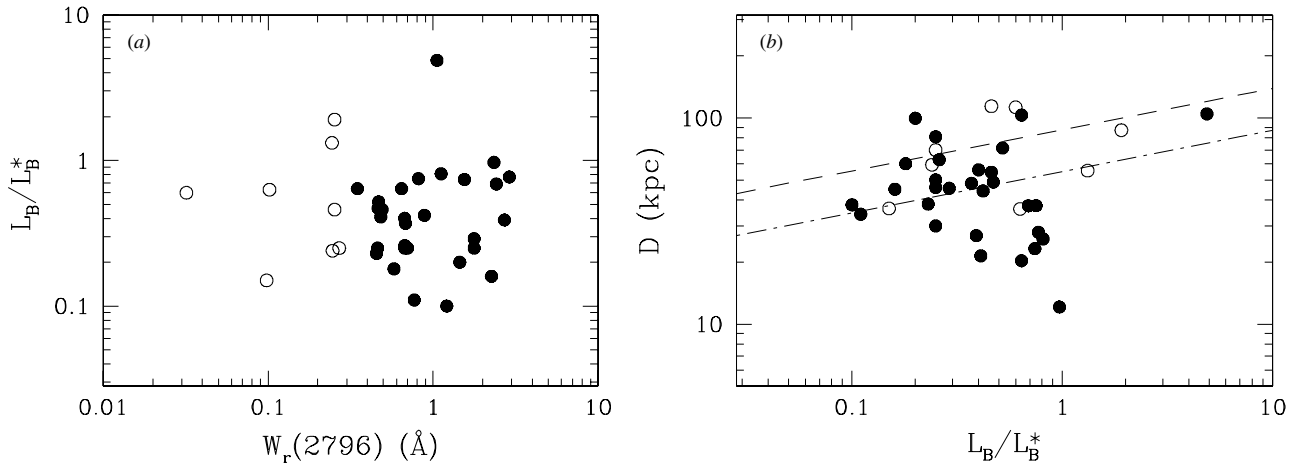


Figure 1. (a) $W_r(2796)$ versus L_B/L_B^* . Filled circles have $W_r(2796) \geq 0.3$ Å and the open circles have $W_r(2796) < 0.3$ Å. (b) The impact parameter, D , versus L_B/L_B^* . The dash-dot line is the halo luminosity scaling given by Equation (1) for the results of S95 ($R_* = 55$ kpc, $f_c = 1$, $\beta = 0.2$). The dash-dash line is the halo luminosity scaling given by Equation (1) for our result ($R_* = 88$ kpc), assuming $f_c = 1$, $\beta = 0.2$.

consider these “weak” systems as “non-absorbing” galaxies in order to be consistent with our comparisons for the remainder of this paper. In Figure 1(a) note that both absorbing and non-absorbing galaxies span the same luminosity range.

3. RESULTS

Applying Equation (2), we computed the statistical absorption radius, R_x , for $W_r(2796) \geq 0.3$ Å employing the most current Schechter luminosity function parameters and absorber redshift path density. We adopted $dN/dz = 0.8$ (Nestor et al. 2005), $\alpha = 1.3$, and $\Phi_* = 3.14 \times 10^{-3}$ Gal Mpc $^{-3}$ (Faber et al. 2007) for the $\langle z \rangle = 0.5$ redshift bin, where the mean redshift of our sample is 0.58. Since the luminosity scaling is not necessarily constrained by our sample, we consider both $\beta = 0.2$ and $\beta = 0$ (i.e., no scaling) for $y = 0.05$ and $y = 0.01$. We obtained

$$R_x = \sqrt{f_c} R_* = \begin{cases} 64 \text{ kpc}, & y = 0.05, \beta = 0 \\ 43 \text{ kpc}, & y = 0.01, \beta = 0 \\ 88 \text{ kpc}, & y = 0.05, \beta = 0.2 \\ 72 \text{ kpc}, & y = 0.01, \beta = 0.2. \end{cases} \quad (3)$$

By *direct fitting* of his sample, S95 empirically deduced $R_* = 55$ kpc and inferred $f_c = 1$, $\beta = 0.2$, and $y = 0.05$. Assuming $f_c = 1$, $\beta = 0.2$, and $y = 0.05$, we computed a statistical covering fraction corrected absorber halo radius of $R_* = 88$ kpc. The difference between the two values arises from the different methods used to determine R_* : S95 applied a fit to a known sample of Mg II absorption selected galaxies, whereas our values are directly computed from measured absorption and galaxy statistics. Assuming f_c less than unity would increase our computed value of R_* , yielding a value even less consistent with that of S95.

In Figure 1(b), the projected quasar–galaxy separation, D , is plotted versus L_B/L_B^* . The mean impact parameter is $\langle D \rangle = 53.2$ kpc which is close to the S95 halo size. The dash-dot line is the halo radius, $R(L)$, from Equation (1) using $R_* = 55$ kpc, $f_c = 1$, and $\beta = 0.2$ found by S95. Three non-absorbing galaxies reside below the $R(L)$ boundary and five reside above. This is not necessarily inconsistent with S95, who found 2 of 14 non-absorbing galaxies below the $R(L)$ boundary. However, we find 16 $W_r(2796) \geq 0.3$ Å absorbers that are outside the

$R(L)$ boundary by as much as 60 kpc. In the standard halo model, galaxies above the $R(L)$ boundary are expected not to be associated with $W_r(2796) \geq 0.3$ Å absorption. The dash-dash line is the halo radius, $R(L)$, from Equation (1) using the parameters $R_x = 88$ kpc, $f_c = 1$, and $\beta = 0.2$. We find that five of the eight non-absorbing galaxies lie below the $R(L)$ boundary. These five galaxies are expected to be strong absorbing galaxies if they obey the $R(L)$ relation. Also, there are three absorbing galaxies above the $R(L)$ boundary.

From Figure 1(b), it would appear that the value of β is not constrained for the B -band luminosities since non-absorbing galaxies are both above and below $R(L)$ for both the $R_* = 55$ kpc deduced by S95 and our computed size of $R_x = 88$ kpc. Assuming that there is no luminosity scaling, we explore halo cross sections with $\beta = 0$. In Figure 2(a), we plot $W_r(2796)$ versus D . The vertical line is the statistical absorber radius, $R_x = 64$ kpc (where $D/R_x = 1$), for $\beta = 0$ and $y = 0.05$. The top axis gives D/R_x . Galaxies to the left of the line are consistent with the computed statistical absorber radius. Galaxies to the right of the line are inconsistent; if the standard halo model applies these particular galaxies must have halos with $f_c < 1$. We find 5 of 29 galaxies at $D > R_x$. If we assume $y = 0.01$ and $\beta = 0$, we obtain $R_x = 43$ kpc and find 16 galaxies reside at $D > R_x$ and that 4 non-absorbing galaxies are expected to have $W_r(2796) \geq 0.3$ Å absorption. Note that R_x is very sensitive to the choice of the luminosity cutoff when $\beta = 0$. Larger β suppresses the faint-end slope in Equation (2), reducing the cross-sectional contribution of the lowest luminosity galaxies that dominate by number.

In Figure 2(b), we plot $W_r(2796)$ versus $D/R(L)$. The dash-dash line is $D/R(L) = 1$ for $R_x = 88$ kpc, $\beta = 0.2$, and $y = 0.05$. Again, three of 29 galaxies have $D/R(L) > 1$ and five non-absorbing galaxies have $D/R(L) < 1$. If we assume $y = 0.01$ and $\beta = 0.2$, we obtain $R_x = 72$ kpc, which increments the number of galaxies at $D/R(L) > 1$ to 7. If we apply $\beta = 0.28$ from Guillemin & Bergeron (1997), then R_x increases to 100 kpc (for $y = 0.05$) and only three absorbing galaxies lie above the $R(L)$ boundary. The dash-dot line is the S95 result where $D/R(L) = 1$ for $R_* = 55$ kpc. It is clear that there is a significant fraction of absorbing galaxies that is not well represented by the standard halo model of S95, since

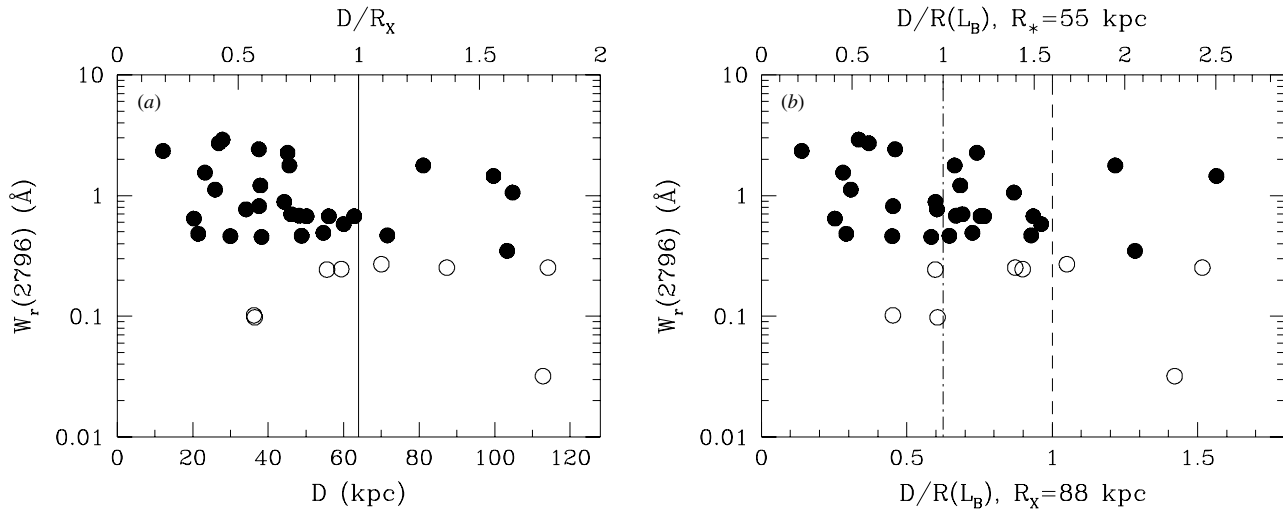


Figure 2. (a) $W_r(2796)$ as a function D (bottom axis) and D/R_x (top axis). The solid line represents $R_x = 64$ kpc for $\beta = 0$ and for a luminosity cutoff of $0.05L_B^*$. (b) $W_r(2796)$ as a function of $D/R(L_B)$. The dashed-dotted line represents $R_* = 55$ kpc and the dash-dash line represents $R_x = 88$ kpc using $\beta = 0.2$ and $y = 0.05$.

Table 1
Mg II Halo Gas Covering Fractions

No. ^a	f_c	y = 0.05		y = 0.01	
		$\beta = 0, R_x = 64$ kpc	$\beta = 0.2, R_x = 88$ kpc	$\beta = 0, R_x = 43$ kpc	$\beta = 0.2, R_x = 72$ kpc
1.	$w_L(\beta, y)(105/R_x)^{-2}$	0.37	0.37	0.17	0.17
2.	$\langle D/R(L_B) \rangle^{-2}$	0.52	0.79	0.58	0.63
3.	$(R_x/55)^{-2}$	0.74	0.40	– ^b	0.58
	$\langle f_c \rangle$	0.54	0.52	0.38	0.46

Notes.

^aThe different methods for computed the covering fractions: (1) the luminosity function weighted f_c , assuming R_* is to be equal to the maximum impact parameter of $D_{\max} = 105$ kpc. (2) The average of the covering fractions for each galaxy was computed for galaxies with impact parameters greater then the statistical halo size. (3) The statistical halo size is assumed to be 55 kpc (S95).

^bOur sample of galaxies provide no constraint on the covering fraction for $y = 0.01$ and $\beta = 0$.

it is expected that spherically symmetric gas halos with unity covering fraction would give rise to absorption exclusively at $D/R(L) \leq 1$.

4. DISCUSSION

Our sample of galaxies is not statistically complete, due to the chosen method of searching for galaxies selected by Mg II absorption (see Churchill et al. 2005). None the less, the data clearly support a covering fraction less than unity, based upon the deduced statistical absorber radius, R_x , which is computed from the redshift path density of the full population of Mg II absorbers.

As seen in Figures 2(a) and (b), a substantial fraction of galaxies are found at impact parameters well beyond the statistical absorber radius. The largest impact parameter in our sample is $D_{\max} \simeq 105$ kpc. If we assume that the largest impact parameter is a proxy for the true size of the covering fraction corrected absorbing halo radius such that $R_* = D_{\max}$, then we can compute a luminosity function weighted covering fraction where $f_c = w_L(\beta, y)(105/R_x)^{-2}$. Assuming a lower galaxy luminosity cutoff of $y = 0.05$, we obtain $f_c = 0.37$ for $\beta = 0$ and $f_c = 0.37$ for $\beta = 0.2$. Assuming a lower luminosity cutoff of $y = 0.01$, we obtain $f_c = 0.17$ for $\beta = 0$ and $f_c = 0.17$ for $\beta = 0.2$. These results are summarized in Table 1. Now with

$f_c < 1$, the presence of eight non-absorbing galaxies within the statistical halo radius, $R(L)$, is consistent with $R_* \simeq 105$ kpc for $W_r(2796) \geq 0.3$ Å.

Using each galaxy from our sample, a conservative estimate of the covering fraction is the mean of the upper limit on $f_c = (D/R(L))^{-2}$. In a complete sample, each galaxy with $D > R(L)$ makes a fractional contribution to reducing the gas covering fraction. Galaxies with $D \leq R(L)$ provide no constraint. If our sample is representative of a complete sample, we obtain $\langle f_c \rangle = 0.52$ ($y = 0.05, \beta = 0$), $\langle f_c \rangle = 0.79$ ($y = 0.05, \beta = 0.2$), $\langle f_c \rangle = 0.58$ ($y = 0.01, \beta = 0$), and $\langle f_c \rangle = 0.63$ ($y = 0.01, \beta = 0.2$).

If we assume that $R_* = 55$ kpc of S95 is the true halo size, we can also compute the covering fractions such that $f_c = (R_x/R_*)^{-2}$. We obtain $f_c = 0.40$ ($y = 0.05, \beta = 0.2$), $f_c = 0.58$, ($y = 0.01, \beta = 0.2$), and $f_c = 0.74$ ($y = 0.05, \beta = 0$). $R_* > R_x$ for $y = 0.01, \beta = 0$ yields no constraint on f_c . The results of the above computations of the covering fractions are summarized in Table 1.

From all the methods of estimating f_c , we obtain $\langle f_c \rangle \sim 0.5$ with a range of $0.17 \leq f_c \leq 0.80$. This is consistent with $f_c = 0.7$ – 0.8 deduced by Charlton & Churchill (1996) from Monte Carlo simulations of Mg II absorption selected galaxy surveys. Our average f_c also consistent with the result of Tripp

& Bowen (2005)⁶ who find $f_c \sim 0.55$ and higher than $f_c \sim 0.25$ determined by Bechtold & Ellingson (1992). Also, Churchill et al. (2007) found a galaxy, probed well within the $R(L)$ boundary, that exhibits no Mg II absorption to $W_r(2796) \leq 7$ mÅ. All these results suggest $f_c < 1$ for Mg II absorbing gas with $W_r(2796) \geq 0.3$ Å. Thus, non-absorbing galaxies below the predicted halo size are expected.

Although the data do not clearly support a halo size–luminosity scaling, if we apply $f_c \sim 0.5$ such that the covering fraction corrected absorbing halo radius is $R_* = 1.41R_x$, a Holmberg-like luminosity relationship with $\beta \simeq 0.2$ is not ruled out for both $y = 0.05$ and $y = 0.01$. We can further constrain f_c , R_* , and β with a maximum likelihood fit that satisfies the distribution of impact parameters and luminosities of our sample. In this analysis, $R_* = \sqrt{f_c}R_x$ is a function of β as constrained by dN/dz . First, we assume that all absorbing galaxies must reside below the $R(L)$ boundary. For $y = 0.05$, we find an upper limit of $f_c \leq 0.4$ for a range of $0.02 \leq \beta \leq 0.24$ with $105 \leq R_* \leq 150$ kpc, respectively. For lower covering fractions, the allowed ranges of β and R_* increase. For $y = 0.01$, we find $f_c \leq 0.2$ for $0.04 \leq \beta \leq 0.66$ with $110 \leq R_* \leq 290$ kpc, respectively. If we relax the condition such that one to three absorption selected galaxies may reside above the $R(L)$ boundary, which could account for errors in the luminosities and/or our finite sample, then the allowed ranges of f_c , R_* , and β increase. For these cases with $y = 0.05$, we find an upper limit of $f_c \leq 0.7$ for $0.18 \leq \beta \leq 0.58$ with $80 \leq R_* \leq 150$ kpc, respectively. Thus, our sample is consistent with a Holmberg-like luminosity relationship in the case $f_c \lesssim 0.5$.

A central issue to this discussion is whether there is a fundamental physical difference between the halos of non-absorbing (weak) and absorbing galaxies or whether the difference in $W_r(2796)$ arises only from a chance intersection of the quasar line of sight through a single gas cloud or a gas cloud complex in these halos. Even if weak $W_r(2796) < 0.3$ Å systems are similar to strong systems, and differ only by the number of clouds intersected along the quasar line of sight, our arguments for constraining the halo gas covering fraction for $W_r(2796) \geq 0.3$ Å still hold.

Strong absorbers are typically characterized by a dominant and blended subsystem and accompanied by significantly weaker subsystems at relative velocities ranging from ~ 40 to 100 km s⁻¹ (Churchill & Vogt 2001). In fact, there may be different physical processes governing the $W_r(2796)$ distribution of weak absorption associated with strong absorbers and the general population of weak absorbers. Churchill et al. (1999) determined that the number density of the general population of weak systems increases as $W_r(2796)$ decreases down to 0.02 Å. In contrast, the Mg II equivalent width distribution of intermediate- and high-velocity subsystems in strong absorbers turns over below $W_r(2796) \sim 0.08$ Å (Churchill & Vogt 2001; Mshar et al. 2007). These facts suggest that lines of sight through galaxy halos often probe a dominant, more massive structure surrounded by smaller fragments of gas; a scenario consistent with patchy halos, in which some lines of sight near galaxies would be expected to probe only $W_r(2796) \gtrsim 0.08$ Å weak absorption.

There is also the possibility that some of the galaxies in our sample having redshifts consistent with those of the Mg II absorbers may not be the primary structure responsible for the absorption. In some cases there could be a faint unidentified galaxy located directly in front of the quasar that cannot be identified even with careful point-spread subtraction of the quasar (see Steidel et al. 1997). Thus, our estimated values of f_c and R_* could be slightly skewed toward smaller and larger values, respectively. It is difficult to quantify this effect since such putative faint galaxies could actually be companions to the galaxies in our sample.

5. CONCLUSIONS

In conclusion, the gas covering fraction must be less than unity since the observed impact parameter distribution of absorbing galaxies does not fall exclusively within the statistical absorber halo radius in the range of $43 \leq R_x \leq 88$ kpc. The fact that some absorbing galaxies are found at $D > R_x$ and some non-absorbing galaxies are found at $D < R_x$ implies $f_c < 1$ and that the standard halo model cannot describe halos on a case-by-case basis. This highlights the power of using the statistics of absorption line surveys to constrain the properties of halos in relation to the measured distributions in absorption selected galaxy surveys.

By quantifying how individual galaxy halos deviate from a “standard” halo, we have obtained an average gas covering fraction of $\langle f_c \rangle \sim 0.5$. It is possible that f_c exhibits both a radial and an equivalent width dependence, though we cannot address this with our sample. Values of f_c are likely to depend on galaxy star formation rates, and galaxy–galaxy mergers and harassment histories; processes that give rise to patchy and geometrically asymmetric gas distributions. Alternatively, the absorption properties of intermediate redshift halos may be governed by the dark matter over density, $\Delta\rho/\rho$, and redshifts at which the galaxies formed (Churchill et al. 2007).

Our results also show that, if $f_c < 1$, the sizes of Mg II absorbing halos can still follow a Holmberg-like luminosity relation with β in the range of 0.2 – 0.28 (S95; Guillemin & Bergeron 1997), which corresponds to $R_* \sim 110$ kpc. If $\beta = 0$ is assumed, then $f_c \leq 0.37$ for our sample to be consistent with no luminosity scaling. In semi-analytical models in which Mg II absorbing gas is infalling and is pressure confined within the cooling radius of hot halos (e.g., Mo & Miralda-Escude 1996; Burkert & Lin 2000; Lin & Murray 2000; Maller & Bullock 2004), a Holmberg-like luminosity relation in quasar absorption line systems naturally arises (Mo & Miralda-Escude 1996). However, these models have great difficulty explaining Mg II absorption at impact parameters greater than ~ 70 kpc. If on the other hand halo gas spatial distributions are governed by stochastic mechanical processes, as suggested by Kacprzak et al. (2007a), then there is no *a priori* reason to expect a clean halo-size luminosity scaling. It is likely that some combination of these scenarios contribute to the statistical values of f_c and β . Thus, it is reasonable to suggest that Mg II halo sizes may not be strictly coupled to the host galaxy luminosity.

Further work on the cross-correlations between absorbers and galaxies would provide better estimates of f_c and β , two quantities that provide direct constraints of galaxy formation simulations. Also needed are additional constraints on the relative kinematics of the absorbing halo gas and galaxies (e.g., Steidel et al. 2002; Ellison et al. 2003; Kacprzak et al. 2007b). What is required is the development of techniques to quantitatively compare observational data with mock quasar

⁶ Tripp & Bowen (2005) have a $W_r(2796) \sim 0.1$ Å (2σ) detection limit which translates to a $W_r(2796) \sim 0.25$ Å (5σ) detection limit as presented here. We have removed all absorbers with $W_r(2796) < 0.25$ Å in order to compare our results at the same detection limit.

absorption line analysis of simulated galaxy halos (Churchill et al. 2006).

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