

HELP: modelling the spectral energy distributions of *Herschel* detected galaxies in the ELAIS N1 field

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ABSTRACT

Aims. The *Herschel* Extragalactic Legacy Project (HELP) focuses on the data from ESA's *Herschel* mission, which covered over 1300 deg² and is preparing to publish a multi-wavelength catalogue of millions of objects. Our main goal is to find the best approach to simultaneously fitting spectral energy distributions (SEDs) of millions of galaxies across a wide redshift range to obtain homogeneous estimates of the main physical parameters of detected infrared (IR) galaxies.

Methods. We perform SED fitting on the ultraviolet(UV)/near-infrared(NIR) to far-infrared(FIR) emission of 42 047 galaxies from the pilot HELP field: ELAIS N1. To do this we use the latest release of CIGALE, a galaxy SED fitting code relying on energy balance, to deliver the main physical parameters such as stellar mass, star formation rate, and dust luminosity. We implement additional quality criteria to the fits by calculating χ^2 values for the stellar and dust part of the spectra independently. These criteria allow us to identify the best fits and to identify peculiar galaxies. We perform the SED fitting of ELAIS N1 galaxies by assuming three different dust attenuation laws separately allowing us to test the impact of the assumed law on estimated physical parameters.

Results. We implemented two additional quality value checks for the SED fitting method based on stellar mass estimation and energy budget. This method allows us to identify possible objects with incorrect matching in the catalogue and peculiar galaxies; we found 351 possible candidates of lensed galaxies using two complementary χ^2 's criteria (stellar and infrared χ^2 's) and photometric redshifts calculated for the IR part of the spectrum only. We find that the attenuation law has an important impact on the stellar mass estimate (on average leading to disparities of a factor of two). We derive the relation between stellar mass estimates obtained by three different attenuation laws and we find the best recipe for our sample. We also make independent estimates of the total dust luminosity parameter from stellar emission by fitting the galaxies with and without IR data separately.

Key words. infrared: galaxies – galaxies: statistics – galaxies: fundamental parameters

1. Introduction

Multi-wavelength data for extragalactic objects is a necessary precondition for a physical analysis of galaxies, as the full complexity of the galaxy is only seen when using different spectral ranges simultaneously. The emission from the hot gas component, active nuclear regions, and the end products of stellar evolution (supernovae and compact remnants) can be observed in the X-rays ($1 < \lambda < 100 \text{ \AA}$; i.e. [Fabbiano 2006](#)). The far-ultraviolet ($\lambda > 912 \text{ \AA}$; FUV) to infrared ($\sim 3 \mu\text{m} - 1 \text{ mm}$ spectral range; IR) spectra of all galaxies arises from stellar light; either directly or reprocessed by the surrounding interstellar medium (ISM). Old stars can be seen in the near-infrared (NIR) spectral range ($0.75 < \lambda < 5 \mu\text{m}$). The dust, composed of a mixture of carbonaceous and amorphous silicate grains ([Draine 2003](#)) is

heated by the interstellar radiation field and emits in the mid- ($\sim 5 < \lambda < 20 \mu\text{m}$) and far-infrared (MIR and FIR, respectively) ($10 \mu\text{m} < \lambda < 1 \text{ mm}$).

Massive young stars dominate the short wavelength range, while evolved stars mainly emit in the NIR ([Kennicutt 1998](#)). The dust component, heated by the interstellar radiation field, can be observed in the wavelength range from NIR to submillimetre (i.e. [Hao et al. 2011](#); [da Cunha et al. 2010](#); [Calzetti et al. 2012](#)). For a more detailed review of the multiwavelength emissions of different galaxy components see the book written by [Boselli \(2011\)](#).

The UV-to-IR spectral energy distribution (SED) contains information about the stars of the galaxy such as the stellar mass or star formation rate (SFR). For example, information on newborn stars can be inferred from the UV data, making the UV

range a very efficient tracer of SFR. Unfortunately, one of the obstacles to observing starburst regions in the UV range only is dust. The newly born stars are surrounded by gas and dust, which, as well as obscuring the most interesting regions, absorb a part of the UV light emitted by stars (e.g. [Buat et al. 2007](#)). Dust grains absorb or scatter photons emitted by stars and re-emit the energy over the full IR range. Therefore, only part of the energy from newly born stars can be observed in UV wavelengths. Infrared emission, reflecting the dust-obscured star formation activity of galaxies ([Genzel & Cesarsky 2000](#)), combined with UV and optical data, can provide a broad range of information about the star formation history (SFH) and SFR.

As shown in previous studies (e.g. [Le Floch et al. 2005](#); [Takeuchi et al. 2005](#)) the fraction of hidden SFR estimated by UV emission increases from ~50% in the Local Universe to ~80% at $z = 1$. [Burgarella et al. \(2013\)](#) combined the measurements of the UV and IR luminosity functions up to redshift 3.6 to calculate the redshift evolution of the total $\text{SFR}_{\text{UV+IR}}$ and dust attenuation. They found that the dust attenuation increases from $z = 0$ to $z = 1.2$ and then decreases to higher redshifts. Therefore the complexity of the SFR seen from UV and IR wavelengths requires the combination of UV and IR data for a proper analysis of the SFH, and its evolution across the history of the universe. Moreover, the ratio between the UV and FIR emission serves as an indicator of the dust attenuation in galaxies (e.g. [Buat 2005](#); [Takeuchi et al. 2005](#)), and the IR wavelengths serve as a tracer for active galactic nuclei (AGNs; e.g. [Leja et al. 2017](#)). All these factors make the multi-wavelength data set a powerful source of information for detailed galaxy studies. The most common and effective approach to obtaining constraints on, for example, stellar masses, SFRs, and dust properties from the large multi-wavelength catalogues of galaxies is to fit physical models to the galaxy's broad-band SED (i.e. [Walcher et al. 2011](#); [Leja et al. 2017](#)). For an overview of recent improvements in models and methods, [Walcher et al. \(2011\)](#) present a detailed review of different techniques for SED fitting. Many new tools were developed for both UV-optical data or the dust part only, such as STARLIGHT ([Cid Fernandes et al. 2005](#)), ULYSS ([Koleva et al. 2009](#)), VESPA ([Tojeiro et al. 2007](#)), Hyperz ([Bolzonella et al. 2000](#)), Le Phare ([Arnouts et al. 1999](#); [Ilbert et al. 2006](#)), PAHFIT ([Smith et al. 2007](#)) or to fit specific types of objects (e.g. AGNfitter, [Calistro Rivera et al. 2016](#)). Also, the Virtual Observatory tools allow for SED fitting. Bayesian analysis is included in many different software packages; for example GOSSIP ([Franzetti et al. 2008](#)), CIGALE ([Noll et al. 2009](#)), and BayeSED ([Han & Han 2014](#)). The wealth of public and private SED fitting tools implies that different surveys tend to be analysed with different tools, with no common set of models or parameters. As a consequence, it is difficult to combine, compare or interpret large datasets for statistical analysis, as different approaches, models, and assumptions result in disparate accuracy, scaling factors, and non-uniform physical parameters across a wide redshift range.

A lack of homogeneous multi-wavelength catalogues (covering over 10% of the entire sky), together with non-uniform physical parameters obtained based on different models and software, makes the analysis of the main physical properties of galaxies statistically limited and biased.

An FP7 project called the *Herschel* Extragalactic Legacy Project (HELP; [Vaccari 2016](#), [Oliver et al. in prep.](#)) funded by European Union will remove the barriers to multi- λ statistical survey science. The main aim of HELP is to provide homogeneously calibrated multi-wavelength catalogues covering roughly 1300 deg² of the *Herschel* Space Observatory

([Pilbratt et al. 2010](#)) survey fields. These catalogues are going to match individual galaxies across broad wavelengths, allowing for multi-wavelength SED fitting to be performed and for statistical studies of the local-to-intermediate redshift galaxy population. The selection criteria, depth maps, and master list creation details will be published together with the catalogues ([Shirley et al. in prep.](#)). The presence of IR data together with UV-NIR counterparts makes the HELP multi-wavelength catalogue a perfect data set for studying galaxy formation and evolution over cosmic time.

In this paper, we present the general HELP strategy for SED fitting for the millions of galaxies observed in multi-wavelength pass-bands. We discuss the software, models, and parameters which are going to be used uniformly for each data set. This approach ensures homogeneity of obtained physical parameters for the final HELP deliverable of 1300 deg² field.

The paper is organised as follows. In Sect. 2 we describe the ELAIS N1 field as a pilot field of the HELP project. In Sect. 3 we present the method applied in this work to fit SEDs and the main physical models and parameters used for the fitting. Section 4 presents all reliability tests as well as two implemented quality checks. In Sect. 5 we discuss the general properties of the sample, while in Sects. 6 and 7 we explore how all physical parameters obtained by fitting SEDs with different dust attenuation models can be biased and whether or not it is possible to predict total dust luminosity from stellar emission only. Our summary and conclusions are then presented in Sect. 8. In this paper we use WMAP7 cosmology ([Komatsu et al. 2011](#)): $\Omega_{\text{m}} = 0.272$, $\Omega_{\Lambda} = 0.728$, $H_0 = 70.4 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

2. Data

In our analysis, we focus on the pilot HELP field: European Large Area ISO Survey North 1 (hereafter ELAIS N1), 9 deg² area centred at 16^h10^m01^s + 54°30'36" ([Oliver et al. 2000](#)). ELAIS N1 is one of 20 fields making up the European Large Area ISO Survey (ELAIS; [Oliver et al. 2000](#); [Rowan-Robinson et al. 2004](#)). The *Herschel* data in ELAIS N1 was obtained as part of the HerMES project [Oliver et al. \(2012\)](#). It is representative of moderately deep fields for future HELP catalogues.

In this paper we briefly summarise the data used for ELAIS N1. A detailed description of the data used for the HELP project (both FIR and ancillary data), the open source pipeline, which was developed for HELP, and the cross-matching procedure, astrometry corrections, and full data diagnostics will be presented in [Shirley et al. \(in prep.\)](#).

2.1. *Herschel* far-infrared sample of Elais N1 field

Herschel was equipped with two continuum imaging instruments, the Photodetector Array Camera and Spectrometer (PACS; [Poglitsch et al. 2010](#)) and the Spectral and Photometric Imaging Receiver (SPIRE; [Griffin et al. 2010](#)). These instruments provided FIR coverage at 100 and 160 μm from PACS, and at 250, 350, and 500 μm from SPIRE. To obtain the photometry of *Herschel* sources a new prior-based source extraction tool was developed, called XID+ ([Hurley et al. 2017](#)).

The XID+¹ is a probabilistic de-blending tool used to extract *Herschel* SPIRE source flux densities from *Herschel* maps that suffer from source confusion ([Nguyen et al. 2010](#)). This

¹ The software is available at https://github.com/H-E-L-P/XID_plus

is achieved by using a Bayesian inference tool to explore the posterior probability distribution. This algorithm is efficient in obtaining Bayesian probability distribution functions (PDFs) for all prior sources, and thus flux uncertainties can be estimated. A detailed description can be found in Hurley et al. (2017). The original XID+, used for ELAIS N1 field, uses information IRAC bands as a prior. We note that XID+ can be run with more sophisticated priors, using both dust luminosity and redshift information (e.g. Pearson et al. 2018, developed a method of incorporating flux predictions from SED fitting procedure as informed priors, finding improvements in the detection of faint sources). The XID+ was run on the ELAIS N1 *Spitzer* MIPS, and *Herschel* PACS and SPIRE map. The flux level, at which the average posterior probability distribution of the source flux becomes Gaussian (see for more details Hurley et al. 2017; indicating that the information from data dominates over the prior), is 20 mJy for MIPS, 12.5 and 17.5 mJy for PACS and for SPIRE is 4 μ Jy for all three bands.

In the deblending work, the priors we use for computing XID+ fluxes must satisfy two criteria: they must have a detection in the *Spitzer* IRAC 1 band and they must have been detected in either optical or NIR (this was done to eliminate artefacts). The entire catalogue of the HELP FIR measurements based on the XID+ tool will be published at the end of the program (2018), together with the full multi-wavelength data collected from other surveys.

2.2. Ancillary data

The HELP master catalogue is built on a positional cross match of all the public survey data available in the optical to MIR range. This comprises observations from the *Isaac Newton* Telescope/Wide Field Camera (INT/WFC) survey (González-Solares et al. 2011), the Subaru Telescope/Hyper Suprime-Cam Strategic Program Catalogues (HSC-SSP; Aihara et al. 2018), the *Spitzer* Adaptation of the Red-sequence Cluster Survey (SpARCS; Tudorica et al. 2017), the UK Infrared Telescope Deep Sky Survey – Deep Extragalactic Survey (UKIDSS–DXS) (Swinbank 2013; Lawrence et al. 2007), the *Spitzer* Extragalactic Representative Volume Survey (SERVS; Mauduit et al. 2012), and the *Spitzer* Wide InfraRed Extragalactic survey (SWIRE; Lonsdale et al. 2003; Surace et al. 2005). We use the *Spitzer* Data Fusion products² for the final two *Spitzer* MIR surveys presented in Vaccari et al. (2010) and Vaccari (2015). The cross match is described in full in Shirley et al. (in prep.). The list of filters is shown in Table 1 and the coverage is presented in Fig. 1.

When multiple measurements are available in similar passbands we take the deepest only since rapidly increasing errors on shallower surveys mean those measurements contribute little to overall photometry. In cases where a detection is available from two similar filters (e.g. MegaCam g and HyperSuprime-Cam g) we check the number of measurements in the full catalogue, the depth associated with each filter, and the distribution of uncertainties. Because there is a large difference in the depths of different surveys, there is no advantage to using multiple measurements where one clearly has an order-of-magnitude lower error and errors in the shallower catalogues may not include systematic errors of comparable size to the random errors. Based on our analysis we decide to use filters in the same order for each band: HyperSuprimeCam – MegaCam – WFC – GPC1. This means that if we have, for example, g band measurements from

² <http://www.mattiavaccari.net/df/>

Table 1. Data used for SED fitting from ELAIS N1 field.

Telescope	Instrument	Filter
CFHT	MegaCam	u, g, r, y, z
Subaru	HSC	$g, r, i, z, N921, y$
<i>Isaac Newton</i>	Wide Field Cam.	u, g, r, i, z
PanSTARRS1	Gigapixel Cam.1	g, r, i, z, y
UKIRT	WFCam	J, K
<i>Spitzer</i>	IRAC	3.6, 4.5, 5.8, 8.0 (μ m)
	MIPS	mips 24 (μ m)
<i>Herschel</i>	PACS	100, 160 (μ m)
	SPIRE	250, 350, 500 (μ m)

WFC and MegaCam, for our analysis, we are going to use only measurements from MegaCam. If we also have HyperSuprime-Cam photometry then only that one would be used for SED fitting.

We also remove objects around bright stars by measuring the size of the circular region around a star that contains no detections as a function of the star magnitude. This compound selection function can be used to model which objects will propagate through to the final SED sample.

2.3. Photometric redshifts

As part of the HELP database, we provide new photometric redshifts generated using a Bayesian combination approach, described in Duncan et al. (2018). They used two independent multiwavelength datasets (NOAO Deep Wide Field Survey Bootes and COSMOS) and performed *zphot* estimation. Duncan et al. (2018) investigated the performance of three *zphot* template sets: (1) default EZY reduced galaxy set (Brammer et al. 2008), (2) “XMM COSMOS” templates (Salvato et al. 2009), and (3) atlas of Galaxy SEDs (Brown et al. 2014) as a function of redshift, radio luminosity, and IR/X-ray properties. They found that only a combination of all template libraries is able to provide a consensus *zphot* estimate and they used a hierarchical model Bayesian combination of the *zphot* estimates.

This method is used for the HELP project and for the next generation of deep radio continuum surveys. The detailed description of the *zphot* methodology and redshift accuracy is presented in Duncan et al. (2018).

2.4. Final sample

The catalogues produced by XID+ contain a flag for those sources that are either below a flux level in which there is a clear detection or the related XID+ Bayesian P-value maps indicate a poor fit in a region local to the source. The final sample includes 50 129 galaxies from the ELAIS N1 survey with flux measurements for PACS or SPIRE data (or both)³.

Our procedure gives us a set of 19 bands: $u, g, r, i, z, N921, y, J, K, Spitzer$ IRAC 3.6, 4.5, 5.8 and 8.0 μ m, *Spitzer* MIPS 24 μ m, and five passbands from *Herschel* PACS (100 and 160 μ m) and SPIRE (250, 350 and 500 μ m). The coverage of the wavelengths is shown in Fig. 1.

³ The full sample and data files can be downloaded at http://hedam.lam.fr/HELP/dataproducts/dmu28/dmu28_ELAIS-N1/data/zphot/ and also accessed with Virtual Observatory standard protocols at <https://herschel-vos.phys.sussex.ac.uk/>.

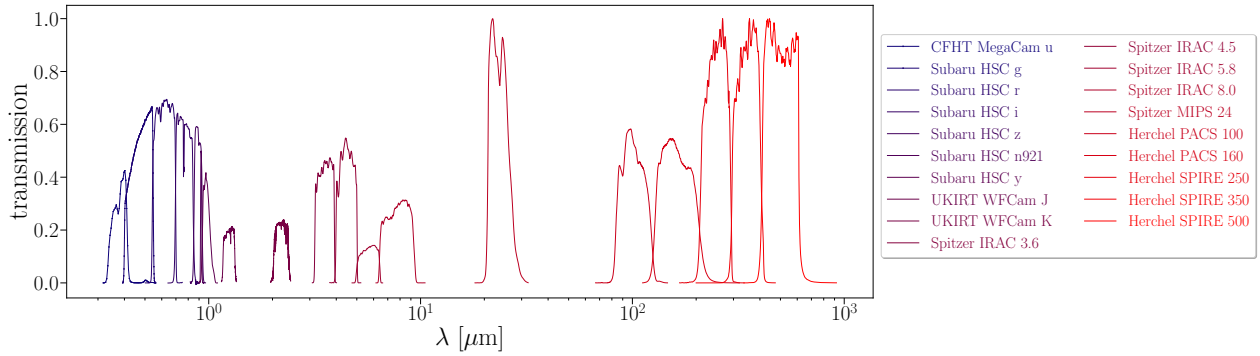


Fig. 1. Transmission curves and demonstrative coverage of primary photometric bands used for SED fitting. Where multiple filters cover the same region, we show only one for clarity. All transmission curves are plotted as exact transmissions.

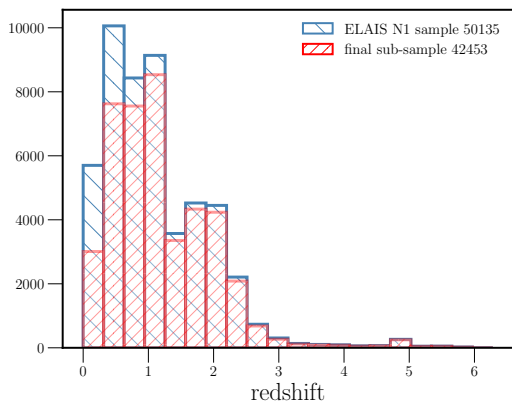


Fig. 2. Redshift distributions of 50 129 galaxies for which it was possible to fit SEDs (blue histogram) and 42 453 galaxies remaining after SED quality cleaning according to Sect. 4.3 (red hatched histogram).

The mean value of the redshift distribution of the final sample is 0.97, while 50% of galaxies (between quartile 1 and 3) are located in the redshift range 0.53–1.63. The redshift distribution of the initial sample can be found in Fig. 2 (blue hatched histogram).

3. Fit of the spectral energy distribution

Taking advantage of the very dense coverage from the broadband passbands (from u band to $500\mu\text{m}$), we use Code Investigating GALaxy Emission⁴ (CIGALE, Boquien et al., in prep.). CIGALE is designed to estimate the physical parameters (i.e. SFR, stellar mass, dust luminosity, dust attenuation, AGN fraction) by comparing modelled galaxy SEDs to observed ones. CIGALE conserves the energy balance between the dust-absorbed stellar emission and its re-emission in the IR. Many authors (e.g. Buat et al. 2015; Ciesla et al. 2015, 2016; Malek et al. 2017) have presented the methodology and the strategy of the code already. A more detailed description of the code will be given by Boquien et al. (in prep.). All adopted parameters used for all modules are presented in Table 2.

3.1. Stellar component

To build the stellar component, we use the stellar population synthesis models by Bruzual & Charlot (2003) with the initial mass function (IMF) given by Chabrier (2003).

⁴ <http://cigale.lam.fr>

As the aim of this analysis is to show how to fit the SEDs of a large sample of galaxies, simplicity and limiting the number of parameters is crucial. It was shown in Ciesla et al. (2015), that a SFH composed of a delayed form to model the bulk of the stellar population with the addition of a flexibility in the recent SFH provides very good estimates of the SFR– M_{star} relation comparing to observations. We refer the reader to Ciesla et al. (2015) for a detailed description of the model. We apply SFH scenarios which include delayed SFR with an additional burst. Parameters which describe our scenario are: age of the galaxy, decreasing rate, burst fraction, and burst age. The functional form for the SFR is calculated as:

$$\text{SFR}(t) = \text{SFR}_{\text{delayd}}(t) + \text{SFR}_{\text{burst}}(t), \quad (1)$$

where $\text{SFR}_{\text{delayd}}(t) \propto te^{-t/\tau_{\text{main}}}$, and $\text{SFR}_{\text{burst}}(t) \propto e^{-(t-t_0)/\tau_{\text{burst}}}$ if $t > t_0$, and $\text{SFR}_{\text{burst}}(t) = 0$ if $t < t_0$. The factor τ_{main} represents e-folding time of the main stellar population and τ_{burst} represents e-folding time of the late starburst population. The adopted parameters used for the stellar component in our fitting procedure are presented in Table 2. We fix values of e-folding times of the main and late stellar population models as those parameters are very difficult to constrain using the SED fitting procedure (i.e. Noll et al. 2009).

3.2. Infrared emission from dust

The current public version of CIGALE (cigale version 0.12.1) includes four different modules to calculate dust properties: Casey (2012), Dale et al. (2007, 2014), and an updated version of the Draine & Li (2007) model, which we decide to use (this model was also used by e.g. Lo Faro et al. 2017 and Pearson et al. 2018 in the framework of the HELP project). In our analysis we apply the multi-parameter dust emission model proposed by Draine et al. (2014), which is described as a mixture of carbonaceous and amorphous silicate grains. Infrared emission SEDs have been calculated for dust grains heated by starlight for various distributions of starlight intensities. The majority of the dust is heated by a radiation field with constant intensity (marked as U_{min} in Table 2) from the diffuse ISM. A much smaller fraction of dust (γ , from Table 2) is illuminated by the starlight with intensity range from U_{min} to U_{max} . This intensity is characterised by a power-law distribution.

We note that we are not able to obtain a reliable value of the γ parameter due to the degeneracy between γ and radiation intensity, and the large photometric uncertainty. We fix the value of γ using the mean value obtained from a stacking analysis

Table 2. Input parameters of the code CIGALE.

Parameter	Values
Delayed star formation history + additional burst	
e-folding time of the main stellar population model (Myr)	3000
e-folding time of the late starburst population model (Myr)	10 000
★ Mass fraction of the late burst population	0.001, 0.010, 0.030, 0.100, 0.300
★ Age (Myr)	1000, 2000, 3000, 4000, 5000, 6500, 10 000
★ Age of the late burst (Myr)	10.0, 40.0, 70.0
Single stellar population: Bruzual & Charlot (2003)	
Initial mass function	Chabrier (2003)
Metallicities (solar metallicity)	0.20
Age of the separation between the young and the old star population (Myr)	10
Attenuation curve	
Main recipe: Charlot & Fall (2000)	
★ A_V in the BCs	0.3, 0.8, 1.2, 1.7, 2.3, 2.8, 3.3, 3.8
Power law slopes of the attenuation in the birth clouds and ISM	-0.7
Lo Faro et al. (2017)	
★ V-band attenuation in the birth clouds (A_V BC)	0.3, 0.8, 1.2, 1.7, 2.3, 2.8, 3.3, 3.8
Power law slopes of the attenuation in the birth clouds and ISM	-0.48
Calzetti et al. (2000)	
★ The colour excess of the stellar continuum light for the young population	0.6, 0.67, 0.74, 0.81, 0.88, 0.95, 1.02, 1.09, 1.16, 1.23, 1.29, 1.36, 1.43, 1.50, 1.57, 1.64, 1.71, 1.78, 1.85, 1.92, 1.99, 2.06, 2.13, 2.2
Dust emission: Draine & Li (2007)	
For our analysis we built templates based on adopted parameters from previous studies: Magdis et al. (2012) , Ciesla et al. (2015) , Lo Faro et al. (2017) , Pearson et al. (2018)	
★ Mass fraction of PAH	1.12, 2.5, 3.19
★ Minimum radiation field (U_{\min})	5.0, 10.0, 25.0
★ Power law slope dU/dM (U^α)	2.0, 2.8
Fraction illuminated from U_{\min} to U_{\max} (γ)	0.02
AGN emission: Fritz et al. (2006)	
For our analysis we used templates built based on average parameters from previous studies: Fritz et al. (2006) , Hatziminaoglou et al. (2009) , Buat et al. (2015) , Ciesla et al. (2015) , Małek et al. (2017)	
Ratio of the maximum to minimum radius of the dust torus	60
★ Optical depth at 9.7 microns	1.0, 6.0
Radial dust distribution in the torus	-0.5
Angular dust distribution in the torus	0.0
Angular opening angle of the torus (deg)	100.0
Angle between equatorial axis and the line of sight (deg)	0.001
★ Fractional contribution of AGN	0.0, 0.15, 0.25, 0.8

Notes. All free parameters are marked with a star before the name.

by [Magdis et al. \(2012\)](#) for the average SEDs of main sequence galaxies at redshifts 1 and 2 ($\gamma = 0.02$).

3.3. AGN component

As shown in the literature (e.g. [Ciesla et al. 2015](#); [Leja et al. 2018](#)) the SED fitting procedure is a powerful technique, but the accuracy of estimated physical properties is tightly correlated with the accuracy of the models used. For example, the AGNs can substantially contribute to the MIR emission of a galaxy.

To improve the derived galaxy properties we add an AGN component to the stellar SED. We derive the fractional contribution of the AGN emission from [Fritz et al. \(2006\)](#) templates, which assume two components: (1) point-like isotropic emission of the central source, and (2) radiation from dust with a toroidal geometry in the vicinity of the central engine. The AGN emission is absorbed by the toroidal obscurer and re-emitted at 1–1000 μm wavelengths or scattered by the same obscurer. We perform SED fitting with a set of parameters from the [Fritz et al. \(2006\)](#) models as described in Table 2. To limit the number of models we fix the value to five variables that parametrizes

the density distribution of the dust within the torus (ratio of the maximum to minimum radius of the dust torus, radial and angular dust distribution in the torus, angular opening angle of the torus, and angle between equatorial axis and the line of sight) with typical values found by [Fritz et al. \(2006\)](#), and used, for example, by [Hatziminaoglou et al. \(2009\)](#), [Buat et al. \(2015\)](#), and [Ciesla et al. \(2015\)](#).

3.4. Dust attenuation

We perform three SED-fitting runs with three different dust-attenuation laws to check which law gives the best fits: we use the [Charlot & Fall \(2000\)](#) model as the main dust attenuation recipe, and then we redo the whole analysis with the popular attenuation law of [Calzetti et al. \(2000\)](#). Subsequently we also test the attenuation law for $z \sim 2$ ULIRGs derived by [Lo Faro et al. \(2017\)](#).

The standard formalism of ([Charlot & Fall 2000](#); hereafter CF00) assumed that stars are formed in interstellar birth clouds (BCs), and after 10^7 yr, young stars disrupt their BCs and migrate into the ambient ISM. The functional form for [Charlot & Fall \(2000\)](#) attenuation law (CIGALE's `dustatt_2powerlaws` module) refers to:

$$A(\lambda) = \begin{cases} A_\lambda(\text{BC}) + A_\lambda(\text{ISM}), & \text{for young stars, age} < 10^7 \text{ yr,} \\ A_\lambda(\text{ISM}), & \text{for old stars (age} > 10^7 \text{ yr).} \end{cases} \quad (2)$$

where $A_\lambda = A_V(\lambda/\lambda_V)^\delta$, λ_V is $550 \mu\text{m}$, A_V is a V-band attenuation, and for CF00 $\delta = -0.7$. The fraction of the total effective optical depth contributed by the diffuse ISM is defined as $\mu = f_{\text{att}}/(1 + f_{\text{att}})$, where $f_{\text{att}} = A_V^{\text{ISM}}/A_V^{\text{BC}}$.

As part of the HELP project, ([Lo Faro et al. 2017](#)) compared power-law attenuation curves with the results of radiative-transfer calculations for 20 ULIRGs observed by *Herschel*. [Lo Faro et al. \(2017\)](#) provided an attenuation law similar to but flatter than that of CF00, which consists of two power laws but with slopes of the attenuation in the BC and ISM equal to -0.48 . The effect of the change of the slopes from -0.7 to -0.48 results in ever greyer attenuation curves for the UV part of the spectrum. The functional form of [Lo Faro et al. \(2017\)](#) attenuation law is exactly the same as Eq. (2) with $\delta = -0.48$.

The functional form for [Calzetti et al. \(2000\)](#) attenuation law is described as:

$$A(\lambda) = E(B - V)k(\lambda), \quad (3)$$

where $k(\lambda)$ is the effective attenuation curve defined as:

$$k(\lambda) = \begin{cases} 2.659(-1.857 + 1.040/\lambda), & \text{for } (0.63 \mu\text{m} \leq \lambda \leq 2.20 \mu\text{m}) \\ 2.659(-2.156 + 1.509/\lambda - 0.198/\lambda^2 + 0.011/\lambda^3), & \text{for } (0.12 \mu\text{m} \leq \lambda < 0.63 \mu\text{m}), \end{cases} \quad (4)$$

and $E(B - V)$ is the colour excess for the stellar continuum.

The difference in shape for attenuation curves defined by [Calzetti et al. \(2000\)](#) and CF00 is mostly present for the wavelengths longer than 5000 \AA , when CF00 curve is much flatter than the one given by [Calzetti et al. \(2000\)](#). A detailed description of the difference between the two attenuation laws can be found in [Lo Faro et al. \(2017\)](#).

To make the comparison possible, we used only one reduction factor between $A_V^{\text{ISM}}/A_V^{\text{BC}}$ (f_{att}) both for CF00 and [Lo Faro et al. \(2017\)](#). As was shown by [Lo Faro et al. \(2017\)](#)

and references therein, the f_{att} parameter is known to usually be unconstrained by broad-band SED fitting. We performed a preliminary analysis using f_{att} as a free parameter (values correspond to μ parameter in the interval 0.2–0.5), and for a final run we chose the most probable value obtained based on the full ELAIS N1 sample: 0.44 ($f_{\text{att}} = 0.8$).

To summarise, the basic attenuation law we used is that of CF00. Then we compared our results with the [Calzetti et al. \(2000\)](#) recipe to check our results with those widely presented in the literature, and then with the [Lo Faro et al. \(2017\)](#) law for ULIRGs. In the next step we performed the comparison between results obtained with all three attenuation laws (see Sect. 6).

3.5. Summary of SED templates

Each created SED template consists of five components (SFH, single stellar population, dust attenuation and emission, and the AGN component) which are spread over the entire grid of input parameters.

All these parameters (with the exception of parameters related to the attenuation curve recipes published by [Calzetti et al. 2000](#) and [Lo Faro et al. 2017](#), which were used for comparing the main physical parameters presented in Sect. 6) are default parameters used for all HELP fields. The full list of the input parameters of the code is presented in Table 2.

Based on the number of parameters presented in Table 2 and the redshift range for our sample of galaxies we built ~ 850 million SED templates which were fitted with CIGALE to 50 129 ELAIS N1 galaxies detected by *Herschel*. We used a 20-core personal computer (PC) with 252Gb of memory. The total time to fit all SEDs for ELAIS N1 was ~ 6 h.

4. Reliability check for SED fitting procedure and identification of outliers

After the SED fitting we removed 143 SED fitting failures (galaxies with χ_r^2 equal to 99⁵) from the initial sample of 50 129 ELAIS N1. From now on we use the remaining 49 986 galaxies as an ELAIS N1 sample.

4.1. Mock analysis

A mock catalogue was generated to check the reliability of the computed physical parameters. To perform this test, we use an option included in CIGALE, which allows for the creation of a mock object for each galaxy for which the physical parameters are known. To build the artificial catalogue we use the best-fit model for each galaxy used for SED fitting (one artificial object per galaxy). A detailed description of the mock analysis can be found in [Lo Faro et al. \(2017\)](#) and [Giovannoli et al. \(2011\)](#). In the next step we perturb the fluxes obtained from the best SEDs according to a Gaussian distribution with σ corresponding to the observed uncertainty for each band. In the final step of the verification of estimated parameters we run CIGALE on the simulated sample using the same set of input parameters as for the original catalogue and compare the output physical parameters of the artificial catalogue with the real ones. A similar reliability check was performed for example by [da Cunha et al. \(2008\)](#),

⁵ The most common reason for SED fitting failures is overestimated redshifts which translates into an estimated age of the galaxy larger than the age of the universe with the assumed ages provided in the CIGALE SED fitting.

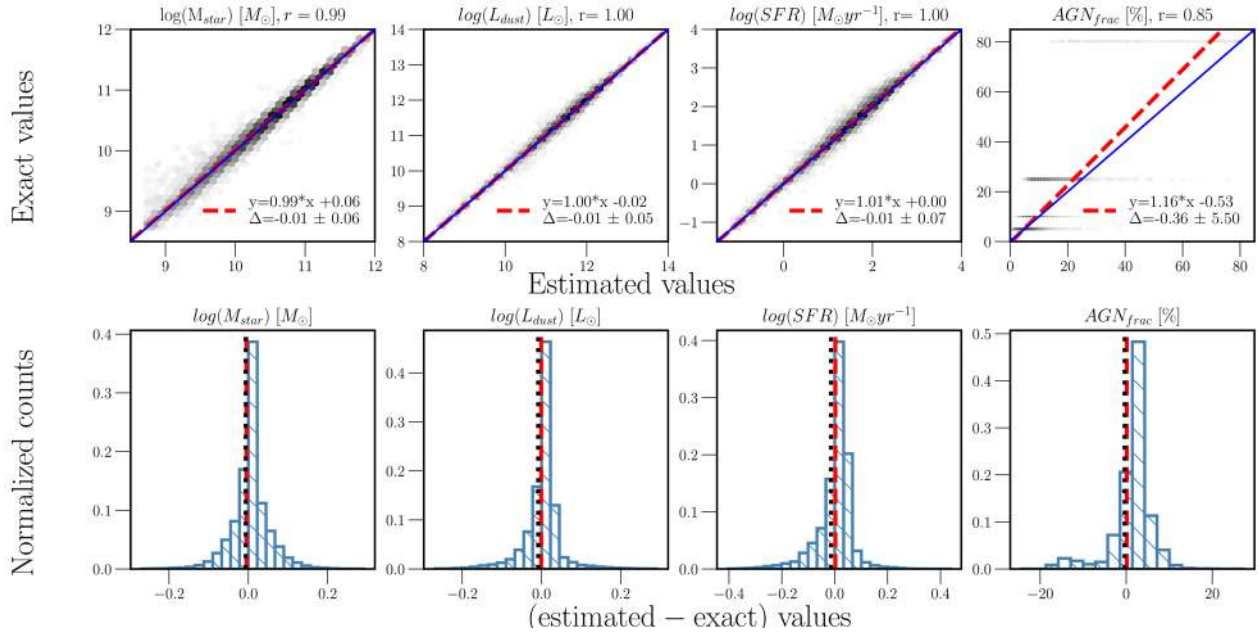


Fig. 3. *Upper panel:* comparison between the true value of the output parameter provided by the best-fit model for the mock catalogue (x-axis) and the value estimated by the code (y-axis), for M_{star} , L_{dust} , SFR and AGN fraction. The Pearson product-moment correlation coefficient is given as an “ r ” value. The blue line corresponds to the 1:1 relation, while the red dashed line is a regression line with the equation given in the legend. Δ represents the mean difference between estimated and exact values and the standard deviation of that difference. *Bottom panel:* distribution of estimated minus exact parameters from upper panel. Red dashed lines correspond to median values, while black dotted lines represent mean values.

Walcher et al. (2011), Yuan et al. (2011), Boquien et al. (2012), Buat et al. (2014), Han & Han (2014) and Ciesla et al. (2015).

All the physical parameters presented in Fig. 3 are computed from their probability distribution function (PDF) as the mean and standard deviation determined from the PDFs (Boquien et al. in prep.), or Walcher et al. (2008) for a more detailed explanation of the PDF method.

The upper panel of Fig. 3 presents the comparison of the output parameters of the mock catalogue with the best values estimated by the code for our real galaxy sample. The value, which characterises the reliability of the obtained properties, is the Pearson product-moment correlation coefficient (r). The dispersion of the main physical parameters is presented in the same plot as the Δ value. The lower panel of Fig. 3 shows the distribution of estimated minus exact values for M_{star} , L_{dust} , SFR, and AGN_{frac} . We find normal distributions with small (~ 0.1) overestimations of $\log(M_{\text{star}})$ and $\log(L_{\text{dust}})$. The third histogram suggests a greater overestimation of SFR which is the result of the slight overestimation of L_{dust} which propagates to SFR. The estimation of the AGN_{frac} is less accurate ($\Delta = -0.36 \pm 5.50$) due to the grid parameters used for a fraction of the AGNs (we were not able to put more than four parameters due to the number of models that needed to be created and the time required to analyse millions of galaxies with all models).

4.2. Quality of the SED fits

The standard global quality of the fitted SEDs is quantified with the reduced value of χ^2 (χ^2 divided by the number of data, hereafter: χ_r^2) of the best model; it is not purely reduced χ^2 regarding the statistical definition as the exact number of free parameters for each galaxy is unknown⁶. The minimum χ_r^2 value

⁶ The number of degrees of freedom can only be estimated for linear models; concerning nonlinear models, the number of degrees of freedom is absolutely nontrivial and whether or not it can be

for the galaxy still points out the best model from the grid of all possible models created with the input parameters, but due to the varied number of observed fluxes and unknown number of free parameters the χ_r^2 criteria cannot be used to remove galaxies with unreliable fits from our sample. Instead of χ_r^2 , we make use of the estimation of physical properties (L_{dust} and M_{star}) to select possible outliers for the SED fitting procedure, and especially, galaxies which do not preserve the energy budget.

To calculate the efficiency of using our method to select the outliers, we choose galaxies for which we can estimate the L_{dust} based on the FIR measurements (the coverage of measurements, Fig. 1, and quality of the optical data allow us to calculate stellar masses for all of the galaxies, and we do not need to make an additional cut for UV–OPT data). To obtain the most reliable sample of galaxies to test our method based on physical properties of galaxies we select objects with at least two measurements for FIR data with signal-to-noise ratios (S/N) ≥ 3 (24 706 galaxies in total, $\sim 50\%$ of the ELAIS N1 catalogue, hereafter: ELAIS N1 3S/N).

We run CIGALE two more times: (1) for optical measurements only, to estimate the classical stellar mass based on UV–OPT measurements and photometric redshifts only (hereafter: $M_{\text{star,stellar SED}}$), and (2) for FIR data only, to calculate dust luminosity (hereafter: $L_{\text{dust,IR SED}}$). In the following steps we compare L_{dust} and M_{star} values obtained from full UV–FIR SED fitting with $L_{\text{dust,IR SED}}$ and $M_{\text{star,stellar SED}}$.

We then select galaxies with dust luminosity and/or stellar masses that are inconsistent with the estimated based on the total SED fitting. We use two criteria to find those objects:

- Criterion 1: L_{dust} inconsistent (within 3σ level) with $L_{\text{dust,IR SED}}$, Fig. 4a.

calculated properly is questionable; see e.g. Andrae et al. (2010) or Chevallard & Charlot (2016).

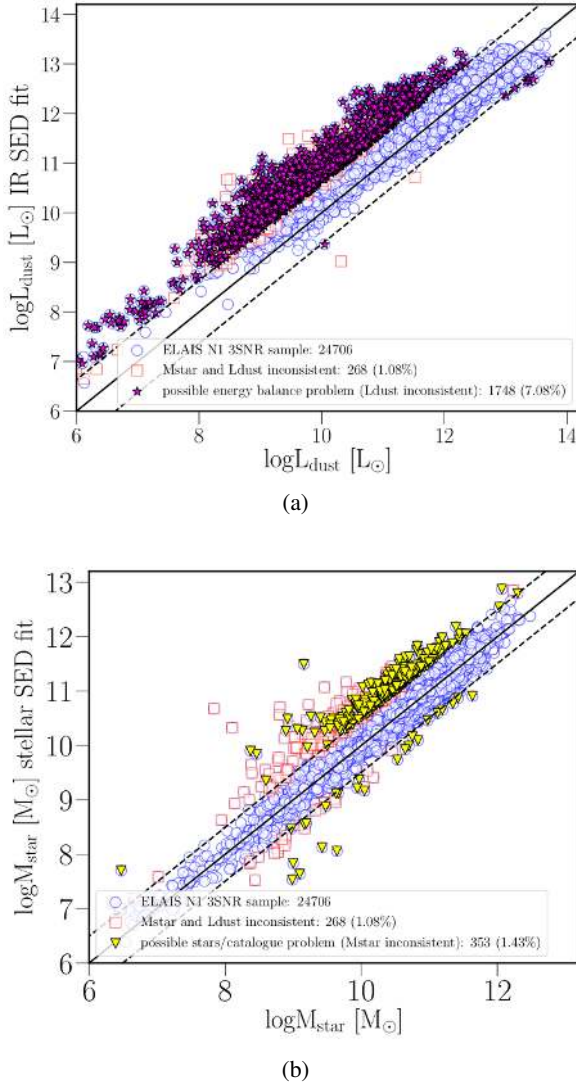


Fig. 4. Comparison between $L_{\text{dust,UV-FIR SED}}$ and $L_{\text{dust,IR SED}}$ (panel a) and $M_{\text{star,UV-FIR SED}}$ and $M_{\text{star,stellar SED}}$ (panel b). The solid black line represents 1:1 relation, while black dashed lines correspond to 3σ . Open blue circles show the ELAIS N1 sample, magenta stars (in panel a) show possible energy budget outliers, filled yellow triangles (panel b) show stellar mass outliers, and open red squares (panels a and b) show objects with inconsistencies in both physical parameters (L_{dust} and M_{star}).

- Criterion 2: M_{star} inconsistent (within 3σ level) with $M_{\text{star,stellar SED}}$, Fig. 4b.

Based on the combination of these two criteria we select three different groups of objects:

- Criterion 1 not criterion 2: energy budget issue; we find 1748 (7.08% of the sample) galaxies with inconsistent L_{dust} and $L_{\text{dust,IR SED}}$ and consistent estimation of stellar mass (full magenta stars in Fig. 4a) and we refer to them as galaxies with possible energy budget issues. Some of the sources are possibly very deeply obscured galaxies (i.e. galaxy in Fig. 5a). After visual inspection we find that $\sim 25\%$ of them also show possible photometric redshift problems (incorrect matching between optical and IR counterparts).
- Criterion 2 not criterion 1: problem with matching optical catalogues; we find 353 objects (1.43% of the sample) with inconsistent M_{star} estimation and at the same time consistent estimation of dust luminosity (objects marked as yellow

triangles in Fig. 4b). Those objects are expected to have problems with matching between catalogues⁷ or there is a problem with calibration between different optical data-sets used for the SED fitting (Fig. 5b).

- Criteria 1 and 2: 268 objects (1.08% of the sample) with inconsistency between runs in both stellar mass and dust luminosity (criteria 1 and 2, red open squares in Figs. 4a and b). The sub-sample is a mixture of stars, incorrect matching between stellar and IR catalogues, possible incorrect photometric redshifts and very small S/N for IR measurements.

In total, based on the physical analysis, we find only 2369 (10% of the ELAIS N1 3S/N sample) peculiar objects (or catalogues mismatches) that should not be included in the further physical analysis.

4.3. Two χ^2 s criteria

To try to remove the objects with inconsistent measurements of L_{dust} (possible energy budget issue) and/or M_{star} (possible problem with matching optical catalogues) labelled in Sect. 4.2, we use a combination of different quantities describing the quality of the fit. Our selection of galaxies can be correlated with χ^2 s calculated for the stellar (hereafter: $\chi_{r,\text{stellar}}^2$) and the dust (hereafter: $\chi_{r,\text{IR}}^2$) parts of the spectra, separately⁸.

Figure 6 shows the scatter between both χ^2 s with groups of outliers found based on the analysis of physical parameters. We define two rejection criteria used to isolate at least 80% of rejected objects as:

$$\chi_{r,\text{IR}}^2 \geq 5.5 \wedge \chi_{r,\text{stellar}}^2 \leq 6, \text{ black solid lines in Fig. 6,} \quad (5)$$

$$\chi_{r,\text{IR}}^2 \geq 11 \wedge \chi_{r,\text{stellar}}^2 \geq 26, \text{ black dashed lines in Fig. 6.}$$

This simple cut based on the values of $\chi_{r,\text{IR}}^2$ and $\chi_{r,\text{stellar}}^2$ allows us to reject:

- 81% of objects with a possible problem with energy balance (magenta stars in Figs. 4a and 6, referred to as group (i) in Sect. 4.2),
- 85% of objects with a possible incorrect match of optical survey or stars (yellow triangles in Figs. 4b and 6, group (ii) in Sect. 4.2),
- and 92% of galaxies with both M_{star} and L_{dust} parameters inconsistent with single runs for stellar and IR parts, respectively (open red squares in Figs. 4a,b and 6, group (iii) in Sect. 4.2).

We stress that the rejection criteria in Eq. (5) are very conservative in order to avoid excluding too many sources; the full sample with all calculated χ^2 s will be released to allow the user to make their own quality cuts. Values used for Eq. (5) are valid for ELAIS N1 field, and the thresholds for other HELP fields should be revised.

Our two χ^2 s criteria (Eq. (5)) remove an additional 2634 galaxies (5.25% of the sample) lying below magenta stars in Fig. 6. We check what kind of objects we reject. We suspect that some of the sources have problems with matching between

⁷ Often with stars – we find that 60% of those objects have the nearest *Gaia* star source from *Gaia* DR1 (*Gaia Collaboration 2016*) between 0.6 and 2 arcsec distance, or that the stellarity parameter is larger than 0.5.

⁸ We defined the wavelength ranges for $\chi_{r,\text{stellar}}^2$ and $\chi_{r,\text{IR}}^2$ as $\leq 8\mu\text{m}$ and $> 8\mu\text{m}$ in rest frame, respectively, and calculate the values directly based on the best model.

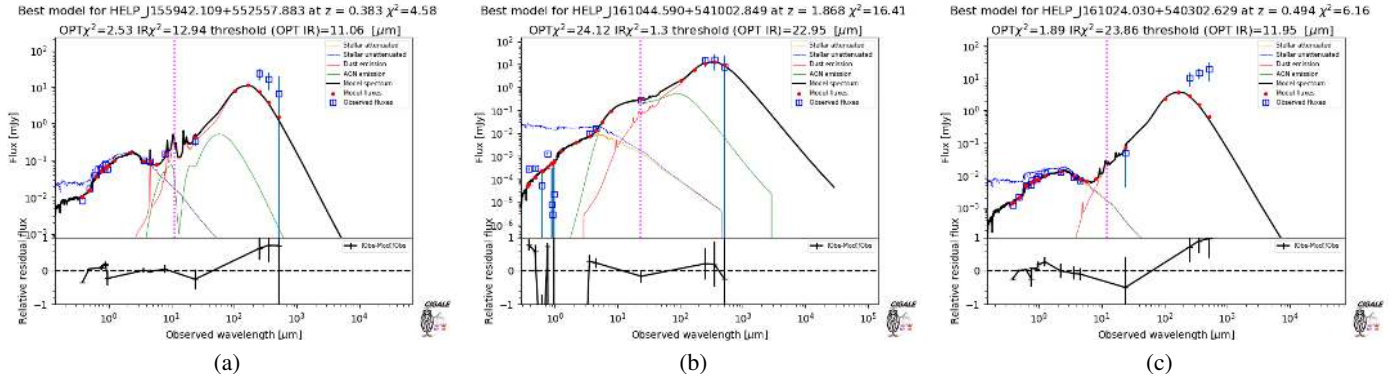


Fig. 5. An example of *panel a*: possible energy budget issue, *panel b*: incorrectly matched optical data, *panel c*: lensed candidate found using a disagreement between stellar and infrared χ^2 of the SED fitting. The optical part of the spectrum is fitted very well when $\chi^2_{r,IR}$ value is on the tail of the $\chi^2_{r,IR}$ distribution. The photometric redshift calculated for the IR part only is equal to 3.28 ± 0.43 , while photometric redshift for the UV–NIR part of the spectrum is 0.49. Observed fluxes are plotted with open blue squares. Filled red circles correspond to the model fluxes. The final model is plotted as a solid black line. The remaining three lines correspond to the stellar, dust, and AGN components. The relative residual fluxes, calculated as $(\text{observed flux} - \text{best model flux})/\text{observed flux}$, are plotted at the bottom of each spectrum.

optical and IR data and that the rejected sample includes possible incorrect photometric redshift, as in, for example, lensed galaxies, which are a good example of objects with energy budget issues.

To check whether or not the mismatch between optical and FIR data is at the basis of the rejection, we calculate $z_{\text{phot,IR}}$ – photometric redshifts based on the PACS and SPIRE data only. We make a standard check for photometric redshift outliers (Ilbert et al. 2006): $\delta(\text{redshift} - z_{\text{phot,IR}})/(1 + \text{redshift}) > N$, where redshift corresponds to the photometric redshift calculated based on UV–OPT data. In the literature, comparisons of photometric and spectroscopic data use a value of N equal to 0.15, but this could be too small for comparison between two photometric values. We find that using double this value ($N = 0.30$, more suitable for less robust estimation) criteria select 78% of rejected objects. Requiring at least three IR measurements with $S/N \geq 3$ for ELAIS N1 catalogues to obtain secure $z_{\text{phot,IR}}$ estimates reduces the ELAIS N1 3S/N sub-sample to 1059 objects only. We also find that for this secure sample, only 10% of objects not rejected by the χ^2 s, and 99% of galaxies selected by Eq. (5), have a possible problem with photometric redshifts. This suggests that the majority of rejected galaxies are incorrect matches between optical and FIR data or peculiar galaxies, as for example lensed galaxies. To summarise, the contamination of rejected objects is as low as $\sim 12\%$.

We checked all objects rejected by Eq. (5) (5291 galaxies in total) and found more than 300 possible lensed candidates inside the criteria describing galaxies with possible energy budget problem (marked in Fig. 6 as black solid line). Based on visual inspection we find that the difference between photometric redshifts estimated based on the UV–OPT and IR measurements for all of them is higher than 0.63 (twice the mean uncertainties obtained for the $z_{\text{phot,IR}}$). An example of a possible lensed candidate is shown in Fig. 5c.

Using two χ^2 s to select peculiar objects is not a new technique (but used for the first time in CIGALE tool). For example, Rowan-Robinson et al. (2014) used the quality of the fit of the IR part of the galaxy spectra to select possible lensing candidates out of 967 SPIRE sources in the HERMES Lockman survey and found 109 candidates. We checked if our method to select lensed candidates (object rejected by Eq. (5) \wedge $|\text{redshift} - z_{\text{phot,IR}}| > 0.63$, the nearest *Gaia* star is further than 0.6 arcsec (we used flag *gaia_flag* == 0 from the HELP

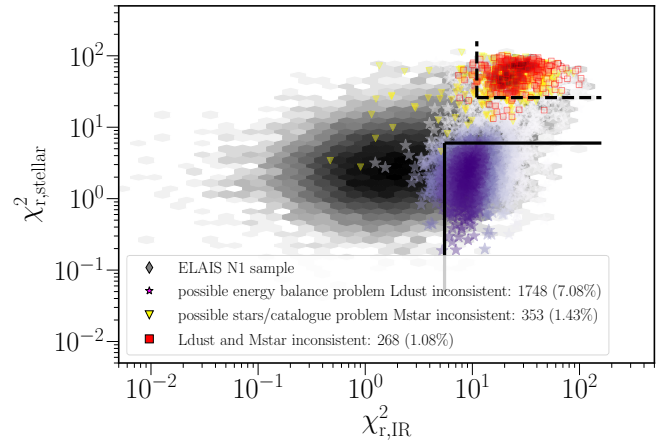


Fig. 6. Scatter between both χ^2 s. Grey-filled hexagons correspond to the full ELAIS N1 sample, magenta stars represent the sample of galaxies with possible energy budget problems, yellow triangles correspond to the galaxies with a possible problem with optical survey matching or stars, and red open squares correspond to objects showing inconsistency for both physical parameters used in our analysis.

catalogue) and stellarity parameter < 0.3) is in agreement with the colour–redshift criteria provided by Rowan-Robinson et al. (2014) for lensed candidates at redshift 0.15–0.95. We find that 70% of our possible lensed galaxies at this redshift range fulfill conditions given by Rowan-Robinson et al. (2014). This proves that this method provides very useful information about the disagreement between the stellar and dust parts of the galaxy spectrum. This information is even more useful for SED fitting, where we are using code that preserves the energy balance.

We also check to see if the galaxies lying between the $\chi^2_{r,stellar}$ rejection criteria (galaxies defined by $5.5 < \chi^2_{r,stellar} < 26$ and $\chi^2_{r,IR} > 11$) have special properties. We find that the majority of them are characterised by very small error bars for UV–NIR measurements and with much higher flux uncertainties for *Herschel* fluxes. This implies imprecise SED fitting, but at the same time, both M_{star} and L_{dust} , calculated separately for UV–IR and IR data, have similar values (below 3σ threshold, with high uncertainties coming from PDF analysis).

Table 3. Main physical properties of the sample of 42 453 galaxies presented in total IR luminosity bins.

(1)	N_{gal} (%)	N_{AGN} (%)	Redshift	$\log(M_{\text{star}})$ (M_{\odot})	$\log(\text{SFR})$ ($M_{\odot} \text{ yr}^{-1}$)	$\log(\text{sSFR})$ (yr^{-1})
(1)	(2)	(3)	(4)	(5)	(6)	(7)
$\log(L_{\text{dust}}) < 11$	6974 16.43%	268 3.86%	0.42 ± 0.15	10.12 ± 0.34	0.64 ± 0.27	-9.70 ± 0.16
LIRGs $11 \leq \log(L_{\text{dust}}) < 12$	23 214 54.68%	724 3.11%	0.96 ± 0.27	10.73 ± 0.24	1.59 ± 0.23	-9.22 ± 0.31
ULIRGs $12 \leq \log(L_{\text{dust}}) < 13$	11 947 28.14%	705 5.90%	1.89 ± 0.32	11.10 ± 0.22	2.30 ± 0.17	-8.85 ± 0.21
HLIRGs $\log(L_{\text{dust}}) \geq 13$	318 0.75%	65 20.44%	4.92 ± 0.34	11.49 ± 0.34	3.52 ± 0.17	-8.01 ± 0.53

Notes. The first column corresponds to the galaxy type according to their total $\log(L_{\text{dust}})$ value, the second shows the number of galaxies (first row), and the total percentage of the sample (second row), Col. (3) presents the number of galaxies with an AGN contribution higher than 20% (first row – total number, second row – percentage value of galaxies which with AGN contribution within the given IR luminosity bin). Median values for redshift, stellar mass, SFR and specific SFR are given in Cols. (4)–(7). The errors are calculated as median absolute deviation.

Nonetheless, we have no solid basis to remove those objects from our analysis.

To provide general properties of the ELAIS N1 sample we removed 7533 galaxies in total based on the two χ^2 s cuts (Eq. (5)). We define the remaining 42 453 galaxies (84% of the original ELAIS N1 sample) as a clean sample for the physical analysis. The redshift distribution of the final sample used for SED fitting is presented in Fig. 2 (the red hatched histogram).

We stress that χ_r^2 , $\chi_{r,\text{stellar}}^2$, and $\chi_{r,\text{IR}}^2$ are a part of deliverables and each user is free to choose different criteria (or to not use them at all). For each field we provide the main physical parameters without any subjective flags describing the quality of the fit. Both, the data and the SED fitting results can be downloaded online⁹ and also accessed with Virtual Observatory standard protocols¹⁰.

5. General properties of the ELAIS N1 sample

5.1. Dust luminosity distribution for ELAIS N1

The majority of our sample consists of (LIRGs) and ultra luminous infrared galaxies (ULIRGs), characterised by $11 \leq \log(L_{\text{dust}}/L_{\odot}) < 12$, and $12 \leq \log(L_{\text{dust}}/L_{\odot}) < 13$, respectively. We have found 23 214 LIRGs (54.68% of the sample), 11 947 ULIRGs (28.14%), and 318 (0.75%) hyper LIRGs (HLIRGs) ($\log(L_{\text{dust}}/L_{\odot}) \geq 13$). The main statistical properties of LIRGs, ULIRGs, and HLIRGs, as well as the galaxies characterised by $\log(L_{\text{dust}}/L_{\odot}) < 11$ from the ELAIS N1 field, can be found in Table 3. All delivered parameters (dust luminosity, stellar mass, SFR, χ_r^2 , $\chi_{r,\text{stellar}}^2$, $\chi_{r,\text{IR}}^2$, and calculated thresholds) will be published while the HELP database is being completed (end of 2018).

5.2. AGN contribution

As it was shown in, for example, Ciesla et al. (2015), Salim et al. (2016) and Leja et al. (2018), that not taking into account an AGN component when performing broad band SED fitting of AGN host galaxies results in substantial biases in their derived parameters.

The AGN fraction in CIGALE is calculated as the AGN contribution to dust luminosity obtained by SED fitting based on

⁹ http://hedam.lam.fr/HELP/dataproducts/dmu28/dmu28_ELAIS-N1/data/zphot/

¹⁰ <https://herschel-vos.phys.sussex.ac.uk/>

Table 4. NIR criteria for an AGN selection.

Criterion	SED AGNs
Stern et al. (2005)	393 (67.7%)
Donley et al. (2012)	339 (58.4%)
Lacy et al. (2007)	435 (75.0%)
Lacy et al. (2013)	485 (83.6%)

Notes. SED AGNs correspond to the number of AGNs found by SED fitting and fulfilling given criterion; numbers in brackets represent the corresponding percentage taking into account the full SED AGN sample of 580 objects with IRAC data.

the bayesian analysis. We find 1762 galaxies (268 galaxies with $\log(L_{\text{dust}}) < 11$, 724 LIRGs, 705 ULIRGs and 65 HLIRG) with a significant ($\geq 20\%$) AGN fraction. We refer to that sample as SED AGNs.

We performed a run of the SED fitting with the same parameters for physical models but without the AGN component to check how the stellar mass and the SFR differ between two runs. We find that $\log(M_{\text{star}})$ is generally well recovered, even for galaxies with a large AGN contribution. Galaxies with a significant AGN contribution have higher SFR when calculating without the AGN component, and the difference can be as high as 0.5 dex. The result of the test is described in Appendix A. We conclude that the use of an AGN component is necessary to run CIGALE's SED fitting for HELP.

We also check whether our method is in agreement with various criteria based on the IRAC colour selections. We select 9196 ELAIS N1 galaxies (21.66% of the final sample) which have measurements at all four IRAC bands. For 580 of them, we find an AGN component according to the CIGALE SED fitting. We focus on four different criteria for AGN selection based on the NIR data: Stern et al. (2005), Donley et al. (2012) and Lacy et al. (2007, 2013). The majority of the 580 SED AGNs meet the listed criteria (see Table 4). Figure 7 shows criteria listed above (areas marked with black dashed lines) and galaxies with a significant AGN contribution according to the SED fitting (marked as red points).

We conclude that the usage of the AGN module (Fritz et al. 2006) for all ELAIS N1 galaxies gives us a reasonable (according to NIR criteria) sample of AGNs. For example, only $\sim 27\%$ (on average) of SED AGNs lie outside the NIR criteria

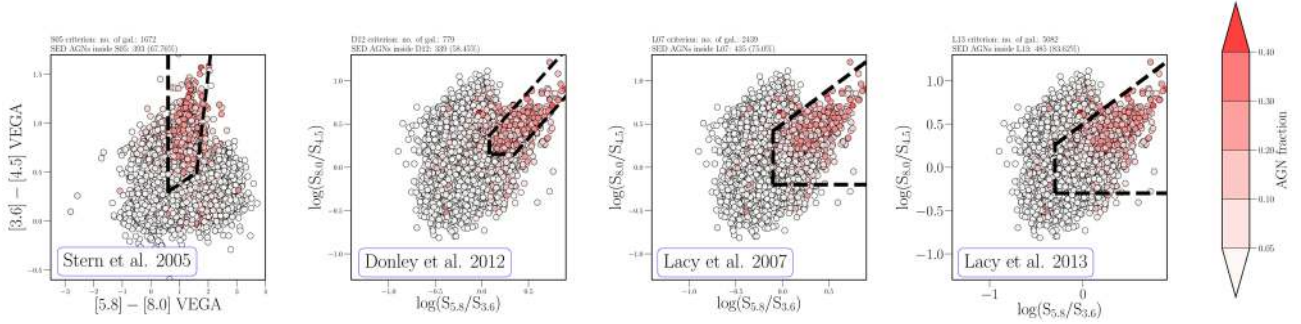


Fig. 7. Stern et al. (2005), Donley et al. (2012) and Lacy et al. (2007, 2013) MIR selection of AGNs. Black dashed lines correspond to the criteria. 546 galaxies from our sample for which the $\text{AGN}_{\text{frac}} \geq 0.2$ are marked in red, according to the AGN fraction.

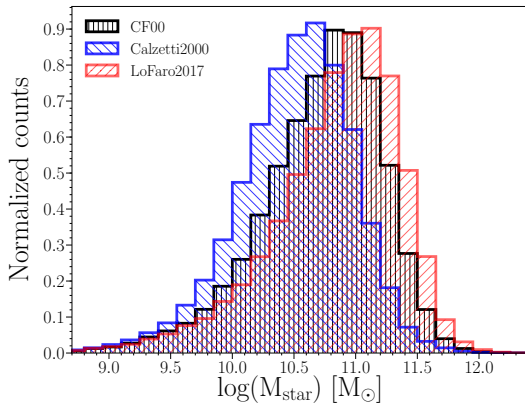


Fig. 8. Comparison of stellar masses obtained with three different dust attenuation laws: Charlot & Fall (2000) – black striped histogram, Calzetti et al. (2000) – blue right hatched histogram, and Lo Faro et al. (2017) – red left hatched histogram.

of Stern et al. (2005), Donley et al. (2012) and Lacy et al. (2007, 2013), meaning that the contamination of the SED AGNs sample is small. In total the ELAIS N1 sample contains only 4.15% of galaxies with a significant AGN contribution and it has no important influence for the main properties of the full sample.

In the framework of the HELP project, we deliver the physical properties of galaxies (M_{star} , SFR, L_{dust}). The AGN identification is a by-product of our analysis. Tests presented above show that the inclusion of emission from dusty AGNs does not generate significant biases for estimated properties.

6. Dust attenuation recipes

Based on our sample we explore how physical parameters obtained by fitting SEDs with different dust attenuation models can be biased. We apply three different attenuation laws to explore this subject. We perform two additional SED fitting runs: (1) the Calzetti et al. (2000) recipe that is widely used in the literature, and (2) the Lo Faro et al. (2017) law obtained for IR bright galaxies. The main change between those attenuation laws is the slope for the part of the spectrum with $\lambda > 5000 \text{ \AA}$ where the CF00 law is much flatter in both the UV and visible part of the spectrum than Calzetti et al. (2000) and a bit steeper than Lo Faro et al. (2017). Figure 1 in Lo Faro et al. (2017) highlights the main differences between all three recipes for dust attenuation considered in this paper.

We find, based on the quality of the fits, that 45% of galaxies are fitted better with CF00 than the other two attenuation

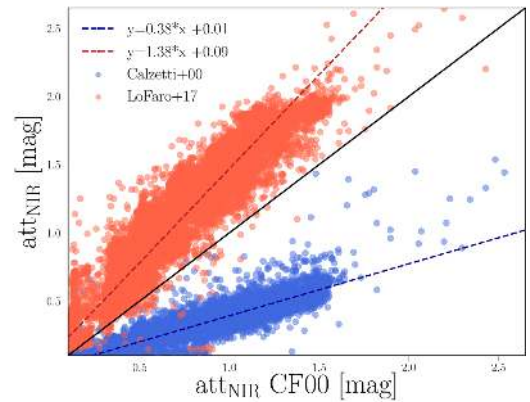


Fig. 9. Total dust attenuation in NIR filter estimated under the assumption of the Charlot & Fall (2000) dust attenuation curve with the Calzetti dust attenuation law (full blue dots) and Lo Faro model (full red dots). The black solid line corresponds to the 1:1 relation while blue and red dashed lines correspond to the relation between Charlot & Fall (2000) and Calzetti et al. (2000), and Charlot & Fall (2000) and Lo Faro et al. (2017) attenuation laws, respectively.

laws. The Lo Faro et al. (2017) procedure works very well for 24% of cases and Calzetti et al. (2000) for 29%. It was expected that CF00 and Lo Faro et al. (2017) attenuation laws would work better for our sample as Lo Faro et al. (2017) has shown that the Calzetti et al. (2000) dust model cannot accurately reproduce the attenuation in a sample of ULIRGs at redshift ~ 2 (around 30% of our sample consists of dusty galaxies at redshift ~ 2 , and more than half of the objects have $\log(L_{\text{dust}}) \geq 11.5 L_{\odot}$).

As was already shown in Lo Faro et al. (2017) and Mitchell et al. (2013), the influence of the attenuation curve on the derived stellar masses can be strong (e.g. Mitchell et al. 2013, showed that for the most extreme cases the difference can reach a factor of 10, with the median value around 1.4) compared to greyer attenuation curves with the standard Calzetti et al. (2000) recipe. A flatter attenuation curve at longer wavelengths results in larger stellar masses. We now compare these findings to our ELAIS-N1 sample. For both additional runs with Calzetti et al. (2000) and Lo Faro et al. (2017) recipes, we used exactly the same parameters for the other modules. All parameters used for SED fitting are listed in Table 2.

For comparison we use only galaxies without possible outliers (we used Eq. (5) presented in Sect. 4.3 for runs with Calzetti et al. (2000) and Lo Faro et al. (2017) attenuation laws). In total we perform a comparison between CF00 and Calzetti et al. (2000) for 37 682 galaxies, and Charlot & Fall (2000) and Lo Faro et al. (2017) for 38 301 objects. We find an

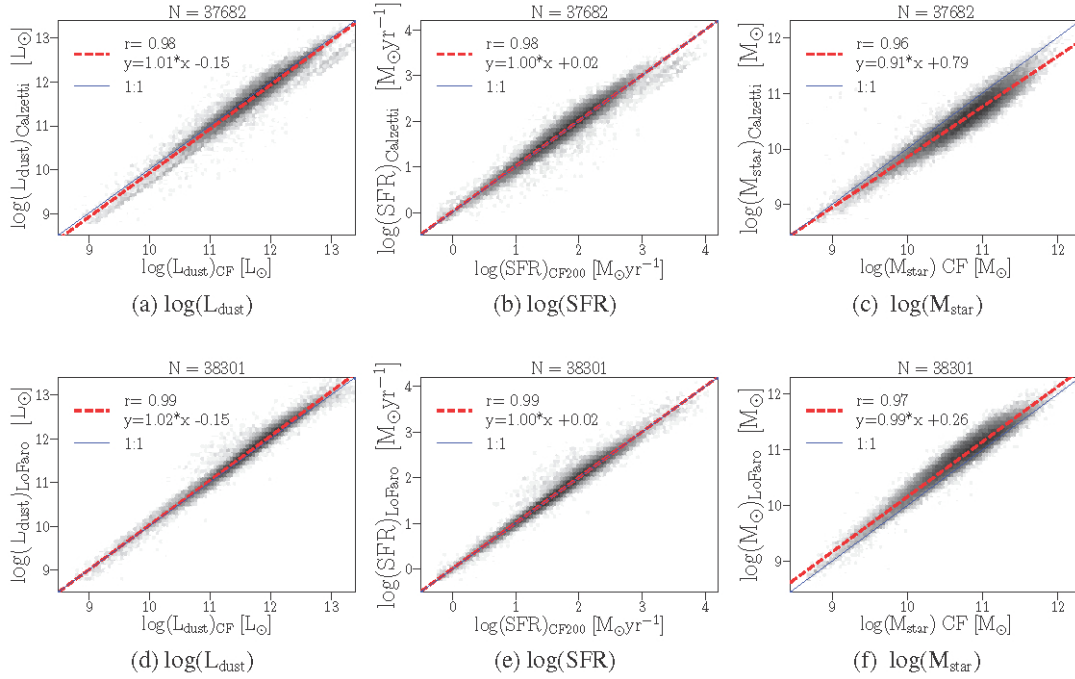


Fig. 10. Comparison of main physical parameters ($\log(L_{\text{dust}})$, $\log(\text{SFR})$, and $\log(M_{\text{star}})$) for a sample of 37 682 galaxies fitted with the Calzetti et al. (2000) attenuation law (y-axis), or the Charlot & Fall (2000) attenuation law (x-axis) (*upper panel*), and for a sample of 38 301 galaxies fitted with the Lo Faro et al. (2017) attenuation law (y-axis), or the Charlot & Fall (2000) attenuation law (x-axis) (*bottom panel*). 1:1 relation is marked with blue solid line, fitted relation – as red dashed line. Each panel includes Pearson product-moment correlation coefficient (r).

agreement between all three runs for L_{dust} and SFR (1:1 with slight scatter: see Fig. 10 for details). We conclude that using either of the attenuation laws of CF00, Calzetti et al. (2000), or Lo Faro et al. (2017) has no important impact on dust luminosity or SFR estimation. Nevertheless, the choice of attenuation law has a substantial impact on calculation of stellar mass.

The right panels of Fig. 10 show the difference between $\log(M_{\text{star}})$ obtained with CF00 and Calzetti et al. (2000) (*upper panel*) and CF00 and Lo Faro et al. (2017) (*bottom panel*) laws. Here we can see a significant shift between both values, which becomes stronger for more massive galaxies for Calzetti versus CF00 attenuation laws.

Figure 8 shows $\log(M_{\text{star}})$ distributions obtained for runs with three different attenuation laws. It is clearly visible that stellar masses obtained with the Calzetti et al. (2000) law are on average lower than those from CF00 and Lo Faro et al. (2017). Derived mean values of $\log(M_{\text{star}})$ [M_⊙] are equal to 10.50 ± 0.47 , 10.72 ± 0.49 , and 10.83 ± 0.54 , for Calzetti et al. (2000), Charlot & Fall (2000), and Lo Faro et al. (2017), respectively.

We find the relation between stellar masses obtained with different attenuation laws: For Calzetti et al. (2000) and CF00:

$$\log(M_{\text{star}})_{\text{Calzetti}} = 0.91 \times \log(M_{\text{star}})_{\text{CF00}} + 0.79, \quad (6)$$

and for Lo Faro et al. (2017) and CF00:

$$\log(M_{\text{star}})_{\text{LoFaro}} = 0.99 \times \log(M_{\text{star}})_{\text{CF00}} + 0.26. \quad (7)$$

We note that we find the same relations excluding IR data from the SED fitting. Our analysis of the influence of the dust attenuation law for estimated dust luminosity, SFR, and stellar mass based on the stellar emission only is presented in Appendix B.

As was also checked by Lo Faro et al. (2017) for a sample of IR selected [U]LIRGs, the total amount of attenuation for

a young population (in UV ranges) is usually well constrained by different attenuation laws, but it is not the case in NIR wavelength range. We find that attenuation in the NIR range is not preserved between different attenuation laws, and is on average larger for Lo Faro et al. (2017) and Charlot & Fall (2000) than for Calzetti et al. (2000; Fig. 9). This relation is a direct outcome of the shape of adopted attenuation curves. We find that the ratio between mean M_{star} obtained from Calzetti et al. (2000) and CF00 is equal to 0.98 ± 0.14 , and that the ratio between mean M_{star} obtained from Lo Faro et al. (2017) and CF00 is 1.49 ± 0.44 . The mean ratio between stellar mass calculated based on the NIR band magnitude from mass to luminosity relation gives very similar values: 0.98 ± 0.01 and 1.41 ± 0.24 for Calzetti/CF00 and LoFaro/CF00, respectively. We summarise that the shape of the attenuation laws in the NIR band is responsible for inconsistency in stellar masses.

The change in stellar mass has an impact on specific SFR and may also introduce a flatter shape for high-mass galaxies at the so-called main sequence galaxies (stellar mass vs. SFR relation, see Fig. 11), as for the high-mass end we have more massive galaxies for Charlot & Fall (2000) and Lo Faro et al. (2017) than for Calzetti et al. (2000) with the same SFR values.

In summary the attenuation law used for SED fitting has an impact on the measure of stellar mass, but has little influence on other parameters, and it should be carefully chosen and taken into account when the properties of stellar masses are discussed.

7. Dust luminosity prediction from stellar emission

The FIR emission is a key component to accurately determine the SFR of galaxies. The strong correlation between L_{dust} and SFR implies that an accurate estimate of total dust luminosity corresponds to a better SFR estimation. However, IR surveys are extremely expensive, and also usually suffer from low

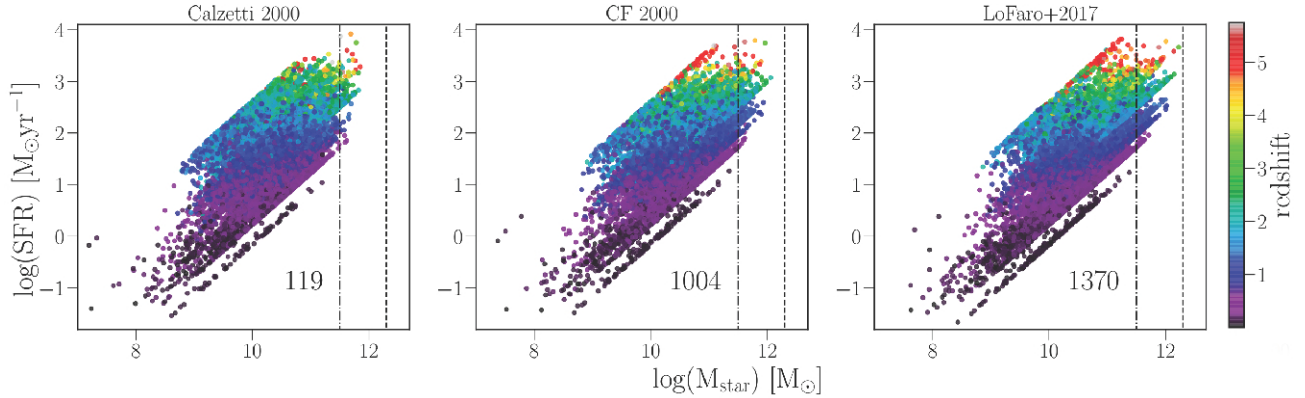


Fig. 11. SFR– M_{star} relation for Calzetti et al. (2000), CF00 and Lo Faro et al. (2017) attenuation laws with colour-coded redshift values. Black dashed-dotted line represents $\log(M_{\text{star}}) = 11.5 M_{\odot}$ and dashed line $\log(M_{\text{star}}) = 12.3 M_{\odot}$. Both lines can help to distinguish the difference between the number of massive galaxies according to the scenario used for SED fitting. Numbers given in each panel correspond to the number of galaxies inside the range $11.5\text{--}12.3 \log(M_{\text{star}})$.

spatial resolution, with respect to optical, and therefore result in source confusion. Moreover, the high-redshift galaxies are mostly explored in the visible and NIR wavelengths only. We investigated the relation between L_{dust} estimation based on different wavelength ranges. Based on the sample of ELAIS N1 galaxies we verified whether or not it was possible to predict the total dust luminosity based on the optical and NIR data only. Two additional runs were performed. We ran CIGALE with the same parameters as for the initial sample of ELAIS N1, but without IR data (rest-frame $\lambda \leq 8 \mu\text{m}$); we refer to that run as *stellar*. We also made a second run with IR data only (*Herschel* PACS and SPIRE measurements), with the same models and parameters (run *IRonly*); we refer to our initial run, with all 19 bands, as *UV–IR*.

For our analysis we use galaxies with at least one PACS measurement with $S/N \geq 2$ and with both SPIRE 250 and SPIRE 350 measurements with $S/N \geq 2$. Selected galaxies have at least six UV–NIR measurements and three *Herschel* measurements (PACS and SPIRE) to cover all strategic parts of the spectra for SED fitting. Our selection, together with χ^2 's criteria (Eq. (5)), results in a sample of 586 galaxies with good coverage of the spectrum and the best photometric measurements.

We check the relation between total dust luminosity for all three runs. First of all we find very good agreement between L_{dust} obtained from *stellar*–IR and *stellar* runs (Fig. 12a). We find the relation to be:

$$\log(L_{\text{dust}})_{\text{stellar}} = (1.00 \pm 0.01) \times \log(L_{\text{dust}})_{\text{UV-IR}} + 0.05. \quad (8)$$

We combined a full run on (UV–IR) with one on *IRonly*, in which dust luminosity was calculated based on the *Herschel* data only. Again, we found very good agreement (see Fig. 12b).

$$\log(L_{\text{dust}})_{\text{IRonly}} = (0.96 \pm 0.01) \times \log(L_{\text{dust}})_{\text{UV-IR}} + 0.46. \quad (9)$$

Finally we performed the comparison of total dust luminosity calculated based on the stellar (*stellar*) data with L_{dust} based on the IR data only (*IRonly*) (Fig. 12c). We have found a clear relation for L_{dust} values:

$$\log(L_{\text{dust}})_{\text{IRonly}} = (0.89 \pm 0.01) \times \log(L_{\text{dust}})_{\text{stellar}} + 1.24. \quad (10)$$

The relation has a much larger slope and scatter but again, we conclude that with CIGALE and the sets of parameters presented in Table 2 we are able to predict the dust luminosity of star forming galaxies based on the UV–NIR data only. The L_{dust}

estimated based on the UV–NIR data is slightly overestimated for faint objects and underestimated for LIRGS and HLIRGS but Eq. (10), which can be used to calibrate L_{dust} obtained from the stellar emission only, takes this into account. Estimated and calibrated values of L_{dust} can be used for a proper estimation of SFR.

8. Summary and conclusions

We present a strategy for SED fitting that is applied to the *Herschel* Extragalactic Legacy Project (HELP) which covers roughly 1300 deg^2 of the *Herschel* Space Observatory. We have focused on the ELAIS N1 as a pilot field for the HELP SED fitting pipeline. We show how the quality of the fits and all main physical parameters were obtained. We present automated reliability checks to ensure that the results from the ELAIS-N1 field are of consistently high quality, with measures of the quality for every object. This ensures that the final HELP deliverable can be used for further statistical analysis.

We introduced the two χ^2 's procedure to remove incorrect SED fits. We calculate two new χ^2 's, apart from a standard χ^2 calculated in CIGALE for the global fit: one for the stellar part ($\chi^2_{r,\text{stellar}}$) and one for the IR part ($\chi^2_{r,\text{IR}}$) of the spectrum. We defined a threshold between $\chi^2_{r,\text{stellar}}$ and $\chi^2_{r,\text{IR}}$ as $8 \mu\text{m}$ in the rest-frame. We demonstrate that the combination of two different χ^2 's can efficiently select outliers and peculiar objects based on the energy balance assumption.

We find that only 4.15% of galaxies of the ELAIS N1 field show an AGN contribution. We check, using four different widely known criteria for AGN selection based on the nNIR data (Stern et al. 2005; Donley et al. 2012; Lacy et al. 2007, 2013), that the majority of AGNs found based on the SED fitting procedure (67.7%, 58.4%, 75.0%, and 83.6% respectively) fulfill the NIR criteria.

We test the influence that different attenuation laws have on the main physical parameters of galaxies. We compare results using Charlot & Fall (2000), Calzetti et al. (2000), and Lo Faro et al. (2017) recipes. We conclude that using different attenuation laws has almost no influence on the calculation of total dust luminosity or SFR, but we find a discrepancy between obtained stellar masses, which we find to be a direct result of the shape of the adopted attenuation curves in NIR wavelengths. We provide relations between stellar masses obtained under those three assumptions of attenuation curves. We find that on

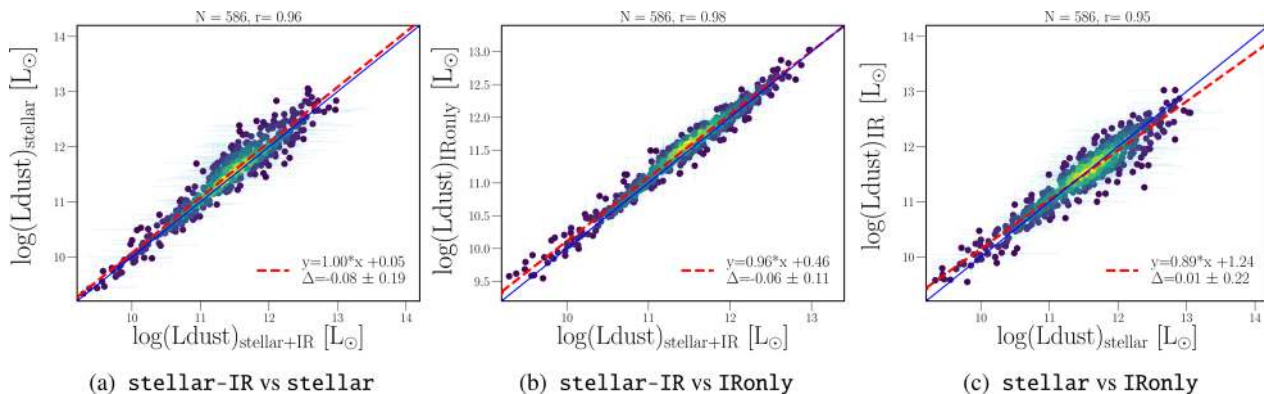


Fig. 12. Comparison of total dust luminosity estimated using different wavelength range for 686 ELAIS N1 galaxies with PACS green and PACS red measurements with $S/N \geq 1.5$, SPIRE 250, and SPIRE 350 $S/N \geq 3$ and at least six UV–NIR measurements; *panel a*: results of SED fitting based on the full sets of data (x -axis) are compared with L_{dust} estimated based on the UV–NIR data only (y -axis); *panel b*: results of SED fitting based on the full sets of data (x -axis) compared to estimates from PACS and SPIRE data only (y -axis); and *panel c*: results of SED fitting based on the UV–NIR data (x -axis) vs. estimates from PACS and SPIRE data only (y -axis). Blue solid lines represent 1:1 relations, while red dashed lines correspond to linear fits to the data. Pearson product-moment correlation coefficient is given as an r value. Δ represents the mean difference between x -axis and y -axis values and the standard deviation of that difference.

average the values of stellar masses for the ELAIS N1 sample can vary up to a factor of approximately two when calculated with different attenuation laws. We demonstrate that the recipe given by Charlot & Fall (2000) more often (for 45% of ELAIS N1 galaxies) outperforms that proposed by Calzetti et al. (2000) and Lo Faro et al. (2017) (according to $\chi^2_{r,\text{IR}}$ and $\chi^2_{r,\text{stellar}}$ values).

We check the accuracy of estimating dust luminosity from stellar emission only and conclude that with CIGALE and the sets of parameters presented in Table 2 we are able to predict L_{dust} for standard IR galaxies, which preserve energy budget, based on the UV–NIR data only. Predicted L_{dust} is in very good agreement with the dust luminosity estimated based on full spectra and stellar emission only. We are not able to estimate monochromatic fluxes but only the total value of L_{dust} . Our tests show that the SFR, tightly correlated with total dust luminosity, is also properly estimated. Our predictions can be used to design new surveys and as priors for the IR extraction pipeline (e.g. for XID+, Pearson et al. 2018).

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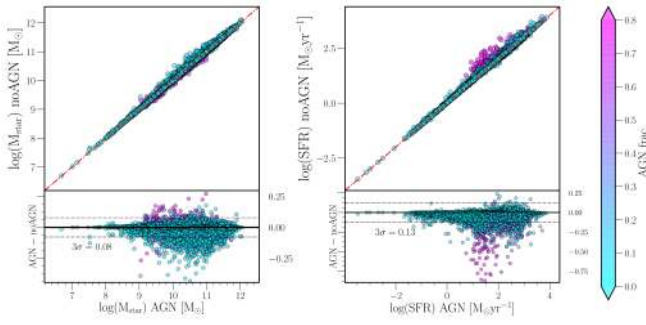


Fig. A.1. Comparison of the $\log(M_{\text{star}})$ and SFR obtained by SED fitting without (y-axis) and with (x-axis) AGN module with colour-coded AGN fraction. Red dashed dotted lines in *upper panels* represent 1:1 relations. *Bottom panels* show difference between both runs. Black solid lines correspond to no difference, while the dashed lines represent 3σ dispersion.

Appendix A: $\log(M_{\text{star}})$ and SFR obtained with and without AGN module

We perform a run of the SED fitting with the same parameters for physical models but without AGN component to check how the stellar mass and the SFR differ between two runs (Fig. A.1). We find that $\log(M_{\text{star}})$ is generally well recovered, even for galaxies with high AGN contribution, with the standard deviation calculated for the difference of $\log(M_{\text{star}})$ estimated with and without AGN module equal to 0.026. Galaxies with significant AGN contribution have higher SFR when calculating without the AGN component. For galaxies characterised by the high fraction of AGN emission the difference can be as high as 0.5 dex. Our result is in agreement with Ciesla et al. (2015) who find that the presence of an AGN can bias SFR estimates starting at AGN contributions higher than 10%, reaching 100% overestimation for AGN fraction $\sim 70\%$ (Ciesla et al. 2015, Appendix A, results for Type 2 AGNs used for our analysis).

Appendix B: Dust attenuation recipes for stellar emission only

We check if the IR data usage can change our result presented in Sect. 6 obtained with data covering wavelength range from optical to IR. We perform the analysis of influence of

Charlot & Fall (2000) and Calzetti et al. (2000) attenuation laws for SED modelling of stellar emission only. From our test we excluded the Lo Faro et al. (2017) formula as an extreme case: this law is much flatter than those by CF00 and Calzetti et al. (2000) and is different suitable for very large attenuation.

We run CIGALE with the same set of modules and parameters as for the main analysis (see Table 2) but without IR data. First we run it with CF00 dust attenuation module and, in the second step, with Calzetti et al. (2000).

We find that SED fitting with Charlot & Fall (2000) law gives better quality fits: 71% of galaxies were fitted better (according to the χ_r^2 values) with CF00 than with Calzetti et al. (2000). This result is in agreement with our previous finding presented in Sect. 6.

Similarly to in Sect. 6, we check the influence of the dust attenuation law on estimations of stellar mass. For comparison we use only galaxies with at least ten optical measurements. In total we perform the comparison for 22 262 galaxies.

The L_{dust} (Fig. B.1a), and consistently the SFR (Fig. B.1b), obtained from two runs are scattered in the region of ULIRGs and HLIRGs (we obtained a similar scatter in Fig. 12c comparing the total dust luminosity estimated using different wavelength ranges. As L_{dust} and SFR are tightly related, the same effect is visible in Fig. B.1b. Figure B.1c shows the relation between $\log(M_{\text{star}})$ obtained with the attenuation laws of Charlot & Fall (2000) and Calzetti et al. (2000) by fitting models for the stellar part only. Obtained relation (Eq. (B.1)) is in a perfect agreement with Eq. (6) obtained for the full spectral fitting (including IR data).

$$\log(M_{\text{star}})_{\text{Calzetti}} = 0.97 \times \log(M_{\text{star}})_{\text{CF}} + 1.02. \quad (\text{B.1})$$

We conclude that the relation between stellar mass estimated with attenuation laws by CF00 and Calzetti et al. (2000) is valid with and without IR data included for SED fitting.

We check whether or not using CF00 or Calzetti et al. (2000) attenuation law to calculate L_{dust} from the stellar emission only can give some bias in function of redshift or L_{dust} estimated based on the full (optical + IR emission) SED fitting. We find that in both cases the scatter is very similar but in case of the Calzetti et al. (2000) recipe we find a clear dependence with L_{dust} and redshift: for bright IR galaxies, the L_{dust} calculated based on the stellar emission only is underestimated. The same relation applies for more distant objects (Fig. B.2).

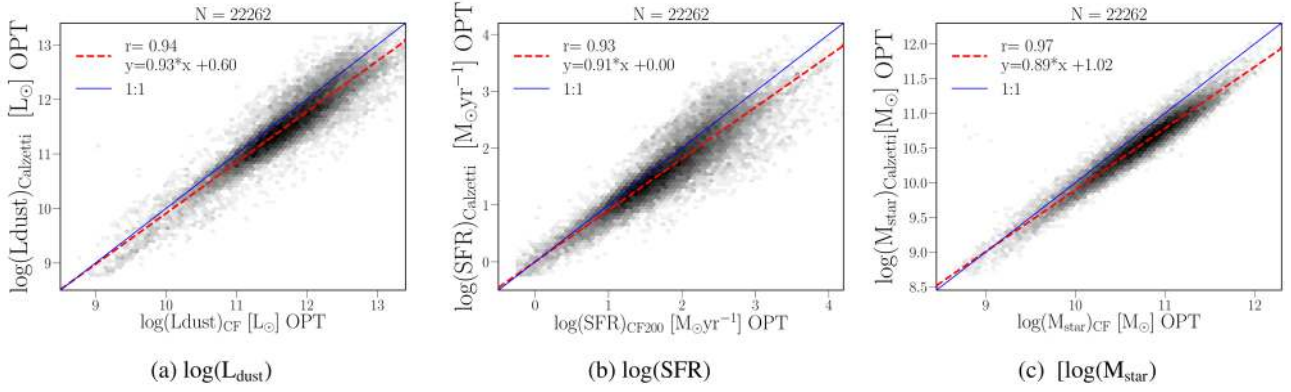


Fig. B.1. Comparison of estimated $\log(L_{\text{dust}})$ (panel a), $\log(\text{SFR})$ (panel b) and $\log(M_{\text{star}})$ (panel c) for a sample of 22262 galaxies fitted with Calzetti et al. (2000) (y-axis) and Charlot & Fall (2000) (x-axis) laws. Blue solid line represents 1:1 relation, while fitted relation is marked as a red dashed line. r represents calculated Pearson product-moment correlation coefficient.

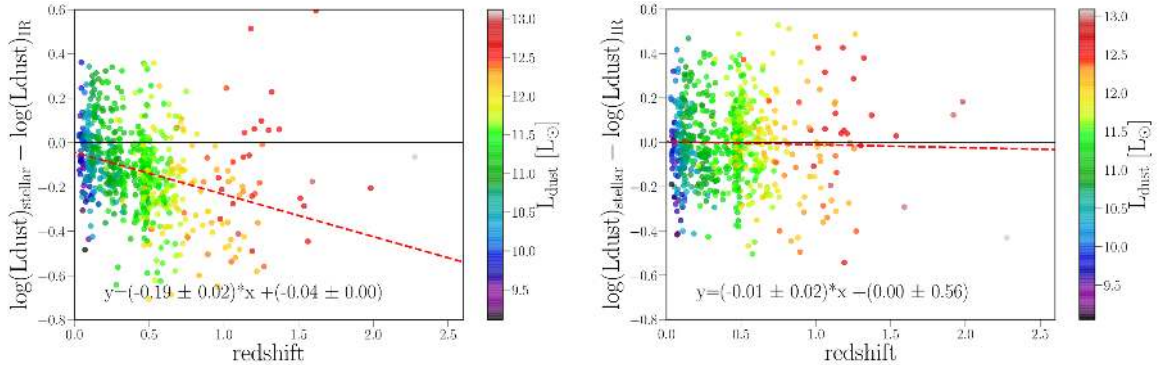


Fig. B.2. The difference between L_{dust} calculated from the stellar emission and IR emission only in function of L_{dust} estimated based on the stellar + IR emission. The left column represents a run with the Calzetti et al. (2000) law while the right column corresponds to the run with the Charlot & Fall (2000) recipe.