

H I and cosmological constraints from intensity mapping, optical and CMB surveys

Alkistis Pourtsidou,[★] David Bacon and Robert Crittenden

Institute of Cosmology and Gravitation, University of Portsmouth, Burnaby Road, Portsmouth PO1 3FX, UK

Accepted 2017 June 12. Received 2017 June 9; in original form 2016 November 20

ABSTRACT

We forecast constraints on neutral hydrogen (H I) and cosmological parameters using near-term intensity mapping surveys with instruments such as BINGO, MeerKAT and the Square Kilometre Array (SKA), and Stages III and IV optical galaxy surveys. If foregrounds and systematic effects can be controlled – a problem that becomes much easier in cross-correlation – these surveys will provide exquisite measurements of the H I density and bias, as well as measurements of the growth of structure, the angular diameter distance and the Hubble rate, over a wide range of redshift. We also investigate the possibility of detecting the late-time integrated Sachs–Wolfe effect using the *Planck* satellite and forthcoming intensity mapping surveys, finding that a large sky survey with Phase 1 of the SKA can achieve a near-optimal detection.

Key words: dark energy – large-scale structure of Universe – cosmology: observations – cosmology: theory.

1 INTRODUCTION

According to the standard cosmological model, dark energy is responsible for the current acceleration of the Universe (Riess et al. 1998; Perlmutter et al. 1999). This is supported by a wealth of high-precision cosmological data, such as measurements of supernovae (e.g. Kowalski et al. 2008; Lampeitl et al. 2009; Betoule et al. 2014), the cosmic microwave background (CMB; Komatsu et al. 2011; Planck Collaboration XIII 2016) and baryon acoustic oscillations (e.g. Blake et al. 2011; Anderson et al. 2014; Aubourg et al. 2015; Delubac et al. 2015). In order to study dark energy’s possible evolution and ultimately uncover its nature, large-scale galaxy surveys like the Dark Energy Survey (DES)¹ and the *Euclid* satellite² are being commissioned. At the same time, a state-of-the-art radio telescope, the Square Kilometre Array (SKA),³ is being built; it aims to answer questions about the formation and evolution of the first galaxies, cosmic magnetism, dark matter, dark energy and gravity.

In recent years, there has been a significant amount of work on the prospects of precision cosmology studies with the SKA and other radio instruments using the neutral hydrogen (H I) intensity mapping (IM) technique (see e.g. Bull et al. 2015; Santos et al. 2015; Pourtsidou et al. 2016). IM (Battye, Davies & Weller 2004; Chang et al. 2008; Loeb & Wyithe 2008; Mao et al. 2008; Peterson

et al. 2009; Seo et al. 2010; Ansari et al. 2012; Battye et al. 2013; Switzer et al. 2013) is a novel technique that uses H I as a dark matter tracer in order to map the 3D large-scale structure (LSS) of the Universe. It measures the intensity of the redshifted 21-cm line, hence it does not need to detect galaxies but treats the 21-cm sky as a diffuse background. This means that IM surveys can scan large areas of the sky very quickly, and perform high-precision clustering measurements.

In order for the aforementioned optical galaxy and IM surveys to reach the unprecedented statistical precision required (<1 per cent for large sky surveys with *Euclid* or SKA), systematic effects must be controlled and removed. In the case of H I IM, for example, we have the presence of large galactic and extragalactic foregrounds. These are orders of magnitude bigger than the signal we want to recover, but they are expected to have a smooth frequency dependence so that they can be removed (Chang et al. 2010; Liu & Tegmark 2011; Wolz et al. 2014, 2015; Alonso et al. 2015; Switzer et al. 2015; Bigot-Sazy et al. 2016; Olivari, Remazeilles & Dickinson 2016). One way to accelerate and improve H I detection and IM cosmology is to cross-correlate the 21-cm maps with optical galaxy surveys, in which case a large part of the noise and systematic effects are expected to be uncorrelated and drop out. In fact, the only H I detections we currently have using IM come from cross-correlating the 21-cm maps from the Green Bank Telescope (GBT) with optical galaxy surveys (Chang et al. 2010; Masui et al. 2013; Switzer et al. 2013). We note that such cross-correlations will benefit optical galaxy surveys as well, since they also suffer from large-scale systematic effects. In general, combining surveys is expected to yield measurements more robust and precise than any individual experiment.

[★] E-mail: alkistis.pourtsidou@port.ac.uk

¹ www.darkenergysurvey.org

² www.euclid-ec.org

³ www.ska.ac.za

In this paper, we will demonstrate how near-term 21-cm IM and optical galaxy surveys can be used to constrain $H\text{I}$ and cosmological parameters. Our work builds upon a previous paper (Pourtsidou et al. 2016), in which we explored the potential of using IM and optical galaxy surveys to detect $H\text{I}$ clustering and weak gravitational lensing and showed that high-precision measurements can be performed. In this work, we will first forecast $H\text{I}$ and cosmological parameter constraints using 21-cm IM autocorrelation clustering measurements, and then move on to cross-correlations with optical galaxy surveys. We will also show how the late integrated Sachs–Wolfe (ISW) effect can be detected by cross-correlating 21-cm IM maps with the CMB maps from the *Planck* satellite mission (Planck Collaboration XIII 2016). As we have already mentioned, the performance of the 21-cm surveys depends on how well foregrounds can be subtracted, as well as on calibration and other problems on very large scales. We have not accounted for these issues in our autocorrelation forecasts, so they should be seen as a ‘best-case’ scenario. Since foreground removal and control of systematics should be much easier in cross-correlation – which is the main focus of this paper – we believe that our cross-correlation forecasts are more realistic and robust. Similar conclusions were reached recently in Wolz et al. (2016), where the authors studied the cross-correlation of IM observations with optical galaxy surveys at $z \approx 0.9$ implementing different star formation models. It was shown that the systematic error contribution to the autopower spectrum measurements can be very large resulting in significant disagreement between theoretical modelling and observations, while cross-correlations perform much better and the modelling is in good agreement with the data from the GBT cross-correlated with WiggleZ galaxies (Masui et al. 2013).

The plan of the paper is as follows: in Section 2, we describe the range of $H\text{I}$ IM and optical galaxy surveys we are going to use for our forecasts. In Section 3, we consider autocorrelation $H\text{I}$ clustering measurements and employ the Fisher matrix formalism to derive constraints for $H\text{I}$ and cosmological parameters. In Section 4, we extend the formalism to include cross-correlations with optical galaxy surveys. Finally, in Section 5 we describe how we can detect the late-time ISW effect in cross-correlation using $H\text{I}$ IM and the *Planck* satellite. We conclude in Section 6.

2 THE SURVEYS

2.1 Intensity mapping surveys

We consider a range of $H\text{I}$ IM surveys, focusing on the SKA and its precursor MeerKAT. As discussed in detail in previous works (Bull et al. 2015), the telescope arrays can be used in autocorrelation (single-dish) mode. This allows the instrument to probe large scales, which is not possible when the array functions in the conventional interferometric mode. We also consider a single-dish experiment, BINGO,⁴ which can be thought of as an excellent pathfinder in order to test and exploit the IM technique, and establish the possible issues that we will face with more advanced IM surveys using MeerKAT and the SKA.

2.2 BINGO

BINGO (Battye et al. 2013) is a single-dish IM project, whose main goal is to detect $H\text{I}$ using the 21-cm IM technique and measure the baryon acoustic oscillation scale in the redshift range $z =$

0.12–0.48 (corresponding to a frequency range of 300 MHz). It will scan approximately 2000 deg^2 in 1 yr total observation time with an angular resolution $\theta_B \simeq 40 \text{ arcmin}$ and $N_{\text{beams}} \sim 50$ feeds. Its system temperature is expected to be $T_{\text{sys}} \simeq 50 \text{ K}$.

The survey’s noise properties have been described in detail in Battye et al. (2013), but we will summarize them here for completeness. The frequency resolution of IM surveys is very good, so we can ignore the instrument response function in the radial direction and only consider the response due to the finite angular resolution:

$$W^2(k) = \exp \left[-k_{\perp}^2 r(z)^2 \left(\frac{\theta_B}{\sqrt{8 \ln 2}} \right)^2 \right], \quad (1)$$

where k_{\perp} is the transverse wavevector, $r(z)$ is the comoving radial distance at redshift z and $\theta_B \sim \lambda/D_{\text{dish}}$ is the beam full width at half-maximum of a single dish with diameter D_{dish} at wavelength λ . Considering a redshift bin with limits z_{min} and z_{max} , the survey volume will be given by

$$V_{\text{sur}} = \Omega_{\text{tot}} \int_{z_{\text{min}}}^{z_{\text{max}}} dz \frac{dV}{dz d\Omega} = \Omega_{\text{tot}} \int_{z_{\text{min}}}^{z_{\text{max}}} dz \frac{cr(z)^2}{H(z)}, \quad (2)$$

and $\Omega_{\text{tot}} = A_{\text{sky}}$, the sky area the survey scans. The pixel volume V_{pix} is also calculated from equation (2), but with

$$\Omega_{\text{pix}} \simeq 1.13 \theta_B^2 \quad (3)$$

assuming a Gaussian beam, and the corresponding pixel z -limits corresponding to the channel width Δf , which is $\sim 1 \text{ MHz}$ for BINGO. Finally, the pixel noise σ_{pix} is given by

$$\sigma_{\text{pix}} = \frac{T_{\text{sys}}}{\sqrt{\Delta f t_{\text{total}} (\Omega_{\text{pix}} / \Omega_{\text{tot}}) N_{\text{dishes}} N_{\text{beams}}}}, \quad (4)$$

with N_{dishes} the number of dishes ($N_{\text{dishes}} = 1$ for BINGO) and t_{total} the total observing time. In IM experiments, the dominant noise contribution comes from the thermal noise of the instrument. The noise power spectrum is then given by

$$P^N(k) = \sigma_{\text{pix}}^2 V_{\text{pix}} W^{-2}(k). \quad (5)$$

2.3 MeerKAT

MeerKAT⁵ is a 64-dish SKA precursor on the planned site of SKA1-MID, with 20 dishes already in place. A detailed description of the noise properties of the MeerKAT dishes can be found in Pourtsidou et al. (2016). MeerKLASS, a large sky IM survey with MeerKAT, has been proposed.⁶ It will scan a few thousand square degrees on the sky (we will take $A_{\text{sky}} = 4000 \text{ deg}^2$ here) in approximately 4000 h total observation time. There are two bands available, namely the L band $0 < z < 0.58$ ($900 < f < 1420 \text{ MHz}$) and the UHF band $0.4 < z < 1.45$ ($580 < f < 1000 \text{ MHz}$); we will show constraints for both bands, taking $z = 0.6$ as the lower z limit for the UHF band. The array will operate in single-dish mode, which means that the autocorrelation signal from the dishes is considered and very large scales can be probed if enough sky is scanned. This means that its noise power spectrum P^N is given by equation (5) with the appropriate parameters for the MeerKAT dishes (Pourtsidou et al. 2016).

⁵ <http://www.ska.ac.za/science-engineering/meerkat/>

⁶ The MeerKLASS survey – Mario Santos, MeerKAT Science: on the Pathway to the SKA, 2016 May 25–27, Stellenbosch, South Africa.

⁴ <http://www.jb.man.ac.uk/research/BINGO/>

2.4 SKA1

We consider Phase 1 of the SKA_MID instrument, consisting of 130 dishes with 15-m diameter according to the recently updated specifications (re-baselining; McPherson 2015). The redshift range is $0.35 < z < 3$ (Band 1) and the expected system temperature is $T_{\text{sys}} = 25$ K. The potential of SKA1 to perform an H I IM survey over a broad range of frequencies and a large fraction of the sky ($f_{\text{sky}} \sim 0.7$) has been presented in Santos et al. (2015). Such a survey can deliver precision measurements of baryon acoustic oscillations, redshift space distortions (RSDs) and weak gravitational lensing, and constrain the spatial curvature of the Universe, primordial non-Gaussianity and the sum of neutrino masses.

2.5 Optical galaxy surveys

Current and future state-of-the-art optical galaxy surveys aim to investigate the nature of dark energy by scanning large areas of the sky and combining multiple cosmological probes, for example, supernovae, baryonic acoustic oscillations, galaxy clusters and weak gravitational lensing. Here, we will consider a Stage III photometric optical galaxy survey, similar to the ongoing DES (see footnote 1), and a future, Stage IV spectroscopic survey, similar to the one that will be performed by the *Euclid* satellite (see footnote 2).

In galaxy surveys, the noise power spectrum is dominated by shot noise, which is due to the discreteness of the galaxies sampling the underlying matter density and is given by

$$P^{\text{shot}} = \frac{1}{\bar{n}} = \frac{1}{(N_g/V_{\text{sur}})}, \quad (6)$$

with N_g the number of galaxies within the redshift bin under consideration with comoving volume V_{sur} .

2.6 Stage III

We consider a Stage III photometric optical galaxy survey with $A_{\text{sky}} = 5000 \text{ deg}^2$, number density of galaxies $n_g = 10 \text{ arcmin}^{-2}$ and redshift range $0 < z < 2$ with median redshift $z_0 = 0.7$. In order to calculate the shot noise contribution for this survey, we will use the above specifications and a redshift distribution of the form (Asorey et al. 2012)

$$\frac{dn}{dz} \propto z^\alpha \exp[-(z/z_0)^\eta], \quad (7)$$

with $\alpha = 2$, $\eta = 3/2$. The photometric redshift error is taken to be $\sigma_z \simeq 0.05(1+z)$. In order to model the galaxy bias on large scales, we will use $b_g = \sqrt{1+z}$ (Rassat et al. 2008).

2.7 Stage IV

We consider a Stage IV spectroscopic optical galaxy survey in the redshift range $0.7 < z < 2$ detecting tens of millions of galaxies in a sky area $A_{\text{sky}} = 15000 \text{ deg}^2$. In our forecasts for such a survey, we will use the number density of galaxies \bar{n} given in Majerotto, Sapone & Schäfer (2016), where the predicted redshift distribution has been split into 14 bins with $\Delta z = 0.1$.

We will now proceed to forecast H I and cosmological constraints from IM autocorrelation measurements, as well as from cross-correlating the 21-cm maps with the optical galaxies.

3 CONSTRAINTS FROM H I INTENSITY MAPPING ALONE

3.1 H I power spectrum

The mean 21-cm emission brightness temperature is given by (see Battye et al. 2013 for a detailed derivation)

$$\bar{T}_b(z) = 180 \Omega_{\text{H I}}(z) h \frac{(1+z)^2}{H(z)/H_0} \text{ mK}, \quad (8)$$

where $\Omega_{\text{H I}}$ is the H I density, $H(z)$ the Hubble parameter as a function of redshift z and $H_0 \equiv 100 h$ its value today. Neglecting for the moment the effect of RSDs, the signal (S) H I power spectrum can be written as

$$P^S \equiv P^{\text{H I}}(k, z) = \bar{T}_b^2 b_{\text{H I}}^2 P(k, z), \quad (9)$$

where P is the matter power spectrum, and $b_{\text{H I}}$ is the H I bias. We can further write

$$P(k, z) = D^2(z) P(k, z=0), \quad (10)$$

where $D(z)$ is the growth factor.

The Fisher matrix for a set of parameters $\{p\}$ is then given by (Tegmark 1997; Seo & Eisenstein 2007)

$$F_{ij} = \frac{1}{2} \left[\mathbf{C}^{-1} \frac{\partial \mathbf{C}}{\partial p_i} \mathbf{C}^{-1} \frac{\partial \mathbf{C}}{\partial p_j} \right] + \frac{\partial \mu^T}{\partial p_i} \mathbf{C}^{-1} \frac{\partial \mu^T}{\partial p_j}, \quad (11)$$

where \mathbf{C} is the covariance matrix and μ is the model for the tested parameters. From the Cramer–Rao inequality, we know that the best errors we can achieve are given by

$$\Delta p_i \geq (F_{ii})^{-1/2}. \quad (12)$$

Following Feldman, Kaiser & Peacock (1994), we can write $\mu_n \approx P^S(k_n)$ in a thin Fourier shell of radius k_n and

$$C_{mn} \approx \frac{2}{V_n V_{\text{sur}}} [P^S(k_n) + 1/\bar{n}]^2 \delta_{mn}, \quad (13)$$

where $V_n \equiv 4\pi k_n^2 dk_n / (2\pi)^3$ is the volume element and dk_n is the width of the shell. Then, we can define the ‘effective volume’ as

$$V_{\text{eff}}(k_n) \equiv \left[\frac{\bar{n} P^S(k_n)}{1 + \bar{n} P^S(k_n)} \right]^2 V_{\text{sur}}, \quad (14)$$

with \bar{n} number density of galaxies and V_{sur} the survey volume. For thick shells that contain many uncorrelated modes, the Fisher Matrix can be written as (Tegmark 1997)

$$F_{ij} \approx \frac{1}{4\pi^2} \int_{k_{\text{min}}}^{k_{\text{max}}} k^2 dk [\partial_i \ln P^S \partial_j \ln P^S] V_{\text{eff}}, \quad (15)$$

where we have also replaced the sum by an integral. For the case of IM surveys, we can write a similar formula (Bull et al. 2015; Poursidou et al. 2016) with $V_{\text{eff}} = V_{\text{sur}} \left(\frac{P^S}{P^S + P^N} \right)^2$, where P^N is the noise power spectrum defined in equation (5). In the following, we will assume that the bias $b_{\text{H I}}$ depends only on the redshift z , i.e. that it is scale independent. This assumption is appropriate only for large (linear) scales, so we will impose a non-linear cut-off at $k_{\text{max}} \simeq 0.14(1+z)^{2/3} \text{ Mpc}^{-1}$ (Smith et al. 2003). Hence, we will also ignore the small-scale velocity dispersion effects (fingers of god). The largest scale the survey can probe corresponds to a wavevector $k_{\text{min}} \simeq 2\pi/V^{1/3}$. Inverting the Fisher matrix, we get the covariance matrix that gives us the forecasted constraints on the chosen parameters set.

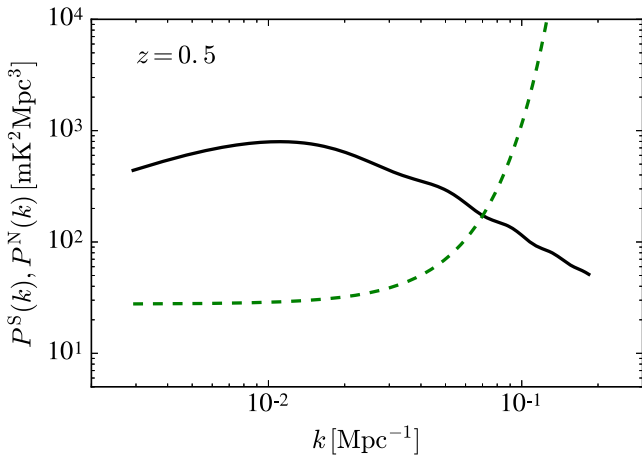


Figure 1. The signal (solid black line) and thermal noise (dashed green line) power spectra at $z = 0.5$ for our fiducial cosmology and the chosen MeerKAT IM survey parameters (see the main text for details).

As a first approach, we assume a flat Λ cold dark matter (Λ CDM) expansion history and keep all cosmological parameters fixed to the Planck 2015 cosmology (Planck Collaboration XIII 2016). Then, the only unknown in the H I power spectrum ($P^{\text{H I}}$) is the prefactor $\Omega_{\text{H I}} b_{\text{H I}}$. Measuring this quantity is very important for future IM surveys, as the magnitude of the power spectrum affects the signal-to-noise ratio (SNR). Pinning down the evolution of the H I density and bias will help us forecast and optimize the scientific output of large sky IM surveys with the SKA.

Knowing how the H I density $\Omega_{\text{H I}}$ evolves with redshift is very important not only for cosmology, but also for galaxy evolution and star formation history studies. At very low redshifts $z < 0.3$, H I galaxy surveys like H I Parkes All-sky Survey can measure $\Omega_{\text{H I}}$ (Zwaan et al. 2005), while at high redshifts $z > 2$ damped Ly α systems can be used (see Crighton et al. 2015 and references therein).

We will now employ the formalism described above to forecast constraints that can be achieved by the MeerKAT IM survey. We calculate the expected constraints on $\Omega_{\text{H I}}(z)b_{\text{H I}}(z)$ across the whole redshift range covered by the L and UHF bands, but we note that the actual survey is going to use only one of the bands. We consider a series of – independent – redshift bins across $0 < z < 1.45$, with width $\Delta z = 0.1$, and use the central redshift of each bin for our calculations. For our fiducial models of the H I density and bias, we use the fits from Bull et al. (2015). As an example of the level of SNR from such a survey, we plot $P^{\text{S}} \equiv P^{\text{H I}}$ and P^{N} for the bin with central redshift $z = 0.5$ in Fig. 1, setting $k_{\perp} \sim k$ in equation (1) for simplicity. We see that the noise diverges as we reach the limits set by the beam resolution – this means that, even if we did not impose a linear cut-off for the k_{max} value in the Fisher matrix, we would not get significant improvement (except at very low z) as the SNR for non-linear scales is dropping fast.

Our results for the fractional uncertainties for $\Omega_{\text{H I}} b_{\text{H I}}$ are presented in Table 1. We see that sub-percent measurements can be achieved at $z < 0.6$ (L band) and 1–4 per cent measurements for $z > 0.6$ (UHF band). These are at least one order of magnitude better than the currently available constraints from galaxy surveys, IM and damped Lyman α observations (see table 2 in Padmanabhan, Choudhury & Refregier 2015). Repeating the same procedure for BINGO – considering one bin with central redshift $z = 0.3$ and $\Delta z \simeq 0.35$ – we find $\delta(\Omega_{\text{H I}} b_{\text{H I}})/(\Omega_{\text{H I}} b_{\text{H I}}) \simeq 0.003$.

Table 1. Forecasted fractional uncertainties on $\Omega_{\text{H I}} b_{\text{H I}}$ and $\Omega_{\text{H I}}$ at various redshifts, assuming the MeerKAT IM survey specifications described in the main text. For the $\Omega_{\text{H I}}$ constraints, we utilize the full H I power spectrum with RSDs.

z	$\delta(\Omega_{\text{H I}} b_{\text{H I}})/(\Omega_{\text{H I}} b_{\text{H I}})$	$\delta\Omega_{\text{H I}}/\Omega_{\text{H I}}$
<i>L</i> band		
0.1	0.010	0.06
0.2	0.005	0.03
0.3	0.005	0.03
0.4	0.007	0.03
0.5	0.009	0.04
UHF band		
0.6	0.011	0.04
0.7	0.013	0.04
0.8	0.015	0.05
0.9	0.018	0.06
1.0	0.022	0.07
1.1	0.026	0.08
1.2	0.030	0.09
1.3	0.036	0.10
1.4	0.042	0.12

3.2 Redshift space distortions

The full H I signal power spectrum in redshift space can be written as

$$P^{\text{S}} \equiv P^{\text{H I}}(k, z; \mu) = \bar{T}_b^2 b_{\text{H I}}^2 [1 + \beta_{\text{H I}}(z)\mu^2]^2 P(k, z), \quad (16)$$

where P is the matter power spectrum, $b_{\text{H I}}$ is the H I bias, $\mu = \hat{k} \cdot \hat{z}$ and $\beta_{\text{H I}}$ is the RSD parameter equal to $f/b_{\text{H I}}$ in linear theory, where $f \equiv \text{dln}D/\text{dln}a$ is the linear growth rate with the scalefactor $a = 1/(1+z)$.

In Masui, McDonald & Pen (2010), the authors showed that RSDs measurements can break the degeneracy between $\Omega_{\text{H I}}$ and $b_{\text{H I}}$. Assuming a flat Λ CDM expansion history, we have $f \simeq \Omega(z)^{0.545}$, and from that we can calculate $D(z)$. Assuming the Planck 2015 cosmology, we parametrize $P^{\text{H I}}$ by two redshift-dependent parameters, namely β and the combination $b_{\text{H I}}^2 \Omega_{\text{H I}}^2$. As discussed in Masui et al. (2010), the amplitude parameter $(b_{\text{H I}} \Omega_{\text{H I}})^2$ will be more precisely measured than β and does not contribute much to the $\Omega_{\text{H I}}$ uncertainty; measuring β we measure $b_{\text{H I}}$, and, subsequently, we get a measurement of $\Omega_{\text{H I}}$:

$$\frac{\delta\Omega_{\text{H I}}}{\Omega_{\text{H I}}} \approx \frac{\delta\beta}{\beta}. \quad (17)$$

The Fisher matrix is now given by (Tegmark 1997; Seo & Eisenstein 2007)

$$F_{ij} = \frac{1}{8\pi^2} \int_{-1}^1 \text{d}\mu \int_{k_{\text{min}}}^{k_{\text{max}}} k^2 \text{d}k [\partial_i \ln P^{\text{S}} \partial_j \ln P^{\text{S}}] V_{\text{eff}}. \quad (18)$$

Our results for the fractional uncertainties for $\Omega_{\text{H I}}$ are presented in Table 1. Note that for the noise modelling we use the full anisotropic response function, equation (1), with $k_{\perp} = k\sqrt{1-\mu^2}$. Again, the results are much better than the currently available (see table 2 in Padmanabhan et al. 2015). In order to demonstrate the potential impact of such measurements, we add our forecasts using the MeerKAT IM survey to the current data coming from H I galaxy surveys and Damped Lyman-Alpha (DLAs) in Fig. 2. As we can see the IM measurements can fill the gap between the low- and high-redshift observations. The combined measurements can be

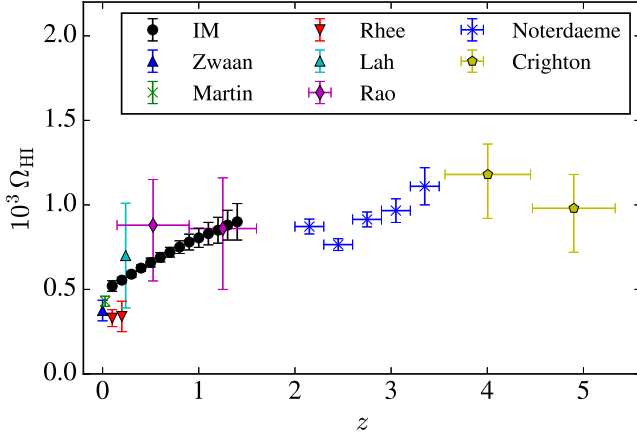


Figure 2. Measurements of the H I density, Ω_{HI} , across redshift. We include the results of Zwaan et al. (2005), Rao, Turnshek & Nestor (2006), Martin et al. (2010), Noterdaeme et al. (2012) and Rhee et al. (2013) [see Crighton et al. (2015) for a complete list], and add our forecasted data for an IM survey with MeerKAT.

used to constrain halo models for cosmological neutral hydrogen (Padmanabhan & Refregier 2017), and to understand the clustering properties of H I by measuring the H I bias (note, for example, that the H I bias and DLAs bias are different, and that IM measures the sum of all H I; Castorina & Villaescusa-Navarro 2016; Sarkar, Bharadwaj & Ananthpindika 2016). It is important to note that these measurements are key in order to discriminate between models of how to assign H I to the dark matter haloes $M_{\text{HI}}(M)$. For example, the simulations of Villaescusa-Navarro et al. (2016) show that AGN feedback reduces M_{HI} and hence suppresses the whole curve, while Padmanabhan & Refregier (2017) suggest a slope effect on the high-mass end of the H I–halo mass relation.

Repeating the above procedure for BINGO (one redshift bin with $0.12 < z < 0.48$), we find $\frac{\delta\Omega_{\text{HI}}}{\Omega_{\text{HI}}} \simeq 0.02$.

We will now modify our approach and assume that the mean temperature $\bar{T}(z)$ has been measured. One way to achieve this is by measuring the smooth part of the H I signal, but global measurements are instrumentally very challenging. However, if one can take advantage of the spectral variation in the brightness temperature, the H I signal can be separated from the foregrounds. This should be doable in Epoch of Reionization measurements, but calibrating the data to the required level is much harder in the low redshifts we consider in this work. Nevertheless, we will consider the H I density known, either from RSD measurements like we described above, or from cross-correlation measurements, H I galaxy surveys and damped Lyman α systems in combination with results from simulations. Alternatively, one could include \bar{T}_b in the RSD functions in the Fisher matrix analysis.

A comprehensive study forecasting cosmological constraints for large sky IM surveys has been presented in Bull et al. (2015). Here, we will first perform a similar analysis in order to see how well a large sky IM survey with MeerKAT can constrain the growth of structure; this will be followed by forecasts using cross-correlations with optical galaxy clustering data. We consider the two approaches (three if we include the measurements coming from optical galaxy surveys alone) highly complementary and synergistic.

Our forecasting method builds upon previous work on geometrical and growth constraints from galaxy redshift surveys like *Euclid* (see e.g. Samushia et al. 2011; Majerotto et al. 2016). Our chosen parameters are the Hubble rate $H(z)$, the angular diameter

Table 2. Forecasted fractional uncertainties on $\{f\sigma_8, D_A, H\}$ at each redshift bin, assuming the MeerKAT IM survey specifications described in the main text.

z	$\delta(f\sigma_8)/(f\sigma_8)$	$\delta D_A/D_A$	$\delta H/H$
<i>L</i> band			
0.1	0.08	0.06	0.07
0.2	0.04	0.03	0.04
0.3	0.03	0.02	0.03
0.4	0.04	0.03	0.03
0.5	0.04	0.03	0.03
<i>UHF</i> band			
0.6	0.05	0.03	0.03
0.7	0.06	0.04	0.03
0.8	0.07	0.05	0.03
0.9	0.08	0.06	0.04
1.0	0.10	0.07	0.04
1.1	0.11	0.08	0.05
1.2	0.13	0.09	0.05
1.3	0.15	0.10	0.06
1.4	0.17	0.12	0.06

$D_A(z)$, $f\sigma_8(z)$ and $b_{\text{HI}}\sigma_8(z)$, which are considered to be independent in each redshift bin. The first two encode important information about the expansion and geometry of the Universe, through their dependence on the matter density Ω_m , the Hubble parameter h , the spatial curvature and the dark energy density Ω_{DE} , while $f\sigma_8(z)$ encodes information about the growth of structure. We choose to work with these parameters because they are model independent. If we want to forecast the constraints on a specific cosmological model, for example, one that can be parametrized by a dark energy equation of state $w = w_0 + w_a z/(1+z)$ (Chevallier & Polarski 2001; Linder 2003) and a growth index γ (Linder & Cahn 2007), we can use a Fisher transformation matrix to move from the old parameters to the new parameters.

The expression of the power spectrum (divided by \bar{T}_b^2) with respect to our chosen parameters is (Seo & Eisenstein 2007)

$$\frac{P^{\text{HI}}(k_f, \mu_f, z)}{\bar{T}_b^2} = \frac{D_A(z)_f^2 H(z)}{D_A(z)^2 H(z)_f} (b_{\text{HI}}\sigma_8(z) + f\sigma_8(z)\mu^2)^2 \frac{P(k, z)}{\sigma_\delta^2(z)}, \quad (19)$$

where the subscript ‘f’ refers to the fiducial (reference) cosmology. Note that the derivatives in the corresponding Fisher matrix are all analytic, except $\partial P(k, z)/\partial k$ that enters the derivatives with respect to D_A and H . That is because $k_{\parallel} = k_{\parallel} H(z)_f/H(z)$, $k_{\perp} = k_{\perp} D_A(z)/D_A(z)_f$.

We perform the Fisher matrix calculations by considering the natural logarithm of the aforementioned parameters. The marginalized constraints on $\{\ln f\sigma_8, \ln D_A, \ln H\}$ for the MeerKAT survey configuration – i.e. the forecasted fractional uncertainties on $\{f\sigma_8, D_A, H\}$ – are presented in Table 2. Here, we also show the covariance matrix corresponding to the full parameter set $\{\ln f\sigma_8, \ln b_{\text{HI}}\sigma_8, \ln D_A, \ln H\}$ for the redshift bin centred at $z = 0.5$:

$$\begin{bmatrix} 0.0019 & -0.00008 & 0.0010 & -0.00004 \\ & 0.0007 & 0.0004 & -0.0007 \\ & & 0.0009 & -0.0004 \\ & & & 0.0007 \end{bmatrix}.$$

From Table 2, we can see that our chosen 4000 deg² survey with MeerKAT can achieve around 4 per cent constraints on $f\sigma_8$ in the L band, i.e. at redshifts where dark energy dominates. Performing the survey in the UHF band worsens the constraints because of the lower SNR of the measurements, but it is a unique opportunity to probe a relatively unexplored cosmological epoch tomographically. The effect of the anisotropic response function of a single-dish IM survey (with excellent resolution along the line of sight but large beam effects along the transverse direction that increase with redshift) is evident – for example, the D_A error increases quickly along the UHF band, while the H error is ‘flatter’. In any case, such a survey is able to provide growth and expansion history constraints for the first time using IM, and complement and compete with state-of-the-art optical galaxy surveys (for a detailed investigation of the possibilities for precision late-time cosmology with 21-cm IM using a variety of radio instruments, see Bull et al. 2015). As we have already stated, the above forecasts are an optimistic scenario, since we have assumed that foregrounds have been removed and we have ignored instrumental systematics and calibration issues at large scales.

A toy model to assess how good the foreground removal has to be in order for IM surveys like the one we are assuming here to yield useful cosmological constraints has been studied in Bull et al. (2015). The contribution from major foregrounds like extragalactic point sources, free–free emission and galactic synchrotron has been parametrized in terms of amplitude and index parameters, and a foreground removal method is assumed. The ‘cleaned’ maps will then have some residual contamination, which is parametrized by an overall amplitude ϵ^2 , with $\epsilon = 1$ corresponding to no removal. There is also a related parameter, namely a minimum radial wavenumber k_{FG} ; the larger that is, the smaller ϵ is (more foreground removal), but obviously there is a tradeoff with loss of cosmological signal. In Bull et al. (2015), it is found that $\epsilon < 10^{-5}$ is needed in order to extract the cosmological information, while $\epsilon = 10^{-6}$ can be achieved using current foreground subtraction techniques.

We will now move on to consider constraints coming from cross-correlating IM maps with optical galaxy data. This is expected to attenuate the above issues.

4 CONSTRAINTS FROM CROSS-CORRELATIONS

In this section, we will investigate what can be achieved from cross-correlating H I IM with optical galaxy surveys. Before we set out the formalism for cross-correlations, it is worth taking a step back and discussing optical galaxy surveys alone. For an optical galaxy survey, the large-scale signal $P^{\text{gg}}(k, z)$ is given by

$$P^{\text{gg}}(k, z) = b_g^2 P(k, z), \quad (20)$$

with b_g the galaxy bias, which we model as linear and deterministic. If we include RSDs, this becomes

$$P^{\text{gg}}(k, \mu, z) = b_g^2 [1 + \beta(z)\mu^2] P(k, z), \quad (21)$$

with $\beta = f/b_g$. The effective survey volume entering the Fisher matrix calculation for the optical galaxy survey autocorrelation measurements can be written as (Seo & Eisenstein 2007)

$$V_{\text{eff}} = V_{\text{sur}} \left(\frac{P^{\text{gg}}(k, \mu, z) \exp[-(k\mu\sigma_r)^2]}{P^{\text{gg}}(k, \mu, z) \exp[-(k\mu\sigma_r)^2] + 1/\bar{n}} \right)^2, \quad (22)$$

with $\sigma_r = c\sigma_z/H(z)$, where σ_z is the redshift error. For a spectroscopic survey like *Euclid*, this is very small (~ 0.001) and can

Table 3. Forecasted fractional uncertainties on $\Omega_{\text{H I}} b_{\text{H I}} r$ at various redshifts, assuming MeerKAT and a Stage IV spectroscopic survey. The sky overlap is taken to be 500 deg², and we have reduced the total observation time (with respect to the full 4000 deg² survey) accordingly.

z	$\delta(\Omega_{\text{H I}} b_{\text{H I}} r) / (\Omega_{\text{H I}} b_{\text{H I}} r)$
UHF band	
0.6	–
0.7	0.06
0.8	0.07
0.9	0.08
1.0	0.10
1.1	0.12
1.2	0.14
1.3	0.16
1.4	0.18

effectively be ignored, but it will be much larger for a photometric survey (~ 0.1). This means that radial information is lost and the power is strongly suppressed along the line-of-sight modes k_{\parallel} (Blake & Bridle 2005; Seo & Eisenstein 2007). For this reason, it is more difficult for photometric surveys to measure observables like $H(z)$ or RSDs.

Let us now set up the formalism for the 21-cm IM – optical cross-correlations. In Masui et al. (2013), the cross-correlation of the 21-cm maps acquired at the GBT and the galaxies of the WiggleZ spectroscopic survey were used to constrain the quantity $\Omega_{\text{H I}} b_{\text{H I}} r$ – with r a correlation coefficient accounting for the possible stochasticity in the galaxy and H I tracers – at $z \sim 0.8$ with a statistical fractional error ~ 16 per cent. The cross-correlation power spectrum can be written as (Chang et al. 2010; Masui et al. 2013)

$$P^{\text{H I, g}}(k, z) = \bar{T}_b b_{\text{H I}} b_g r P(k, z), \quad (23)$$

where we have neglected RSDs (we will include them in the next section). Assuming a flat Λ CDM expansion history keeping all cosmological parameters fixed to the best-fitting Planck cosmology, and using the measurement of b_g by the galaxy survey, as in Masui et al. (2013), the only unknown is the prefactor $\Omega_{\text{H I}} b_{\text{H I}} r$.

Here, we will first consider a forthcoming Stage IV spectroscopic survey (similar to *Euclid*), which is assumed to have a 500 deg² overlap with MeerKAT. This is a conservative assumption since we do not yet know which part of the sky will be scanned by the MeerKAT IM survey. As we have already stated, the *Euclid* spectroscopic redshift range is $0.7 < z < 2$, which overlaps with the UHF band for MeerKAT. Our fiducial model is as before and we set the fiducial stochastic correlation coefficient $r = 1$.

The Fisher matrix in this case will be given by equation (15) with

$$V_{\text{eff}} = V_{\text{sur}} \frac{P^{\text{H I, g}}(k, z)^2}{P^{\text{H I, g}}(k, z)^2 + P^{\text{H I, tot}}(k, z) P^{\text{g, tot}}(k, z)}, \quad (24)$$

with $P^{\text{H I, tot}} = P^{\text{H I}} + P^{\text{N}}$ the total (signal+noise) H I power spectrum, and $P^{\text{g, tot}} = P^{\text{gg}} + 1/\bar{n}$. We take \bar{n} to be similar to the mean number density (in each redshift bin) used in Majerotto et al. (2016). The constraints from a Stage IV spectroscopic survey and MeerKAT with 500 deg² overlap on the sky are summarized in Table 3. We can see that even with a ‘small’ (compared to the sky area of the individual surveys) sky overlap, we can measure $\Omega_{\text{H I}} b_{\text{H I}} r$ across redshift with quite good precision – allowing 4000 deg² overlap we

get a factor of ~ 3 improvement in the constraints. Note that for an ‘intermediate’ spectroscopic survey with the same sky and redshift overlap, but with 100 times smaller number density of galaxies in each bin, we get constraints that are a factor of ~ 2 worse.

As we have already mentioned, the abundance and bias properties of H I are of great importance for both astrophysics (for example, determining the H I luminosity function) and cosmology. The capabilities and potential science output of future H I IM surveys depend on the amplitude of the signal $\Omega_{\text{H I}} b_{\text{H I}}$. Better measurements of $b_{\text{H I}}$ and r are needed, but we can assume that – especially on large scales – the value of r is scale independent and consistent with unity (Khandai et al. 2011; Wolz et al. 2016). Therefore, cross-correlation studies can put a lower limit on $\Omega_{\text{H I}} b_{\text{H I}}$ (Masui et al. 2013).

4.1 Redshift space distortions

Including RSDs in the IM–optical galaxy cross-correlation formalism (see White, Song & Percival 2008 for the case of multiple galaxy populations), we can write

$$P^{\text{H I, g}}(k, \mu, z) = \bar{T}_b b_{\text{H I}} b_g [1 + \beta_{\text{H I}} \mu^2] [1 + \beta_g \mu^2] P(k, z). \quad (25)$$

Note that we have not included stochastic coefficients (r factors), i.e. we have implicitly assumed that they are scale independent and equal to unity on large scales.

We are first going to consider the combination of MeerKAT and a Stage III photometric optical galaxy survey, with a 4000 deg² overlap. We are going to incorporate the photometric redshift error $\sigma_z \simeq 0.05(1+z)$ by replacing (Seo & Eisenstein 2007)

$$P^{\text{H I, g}} \rightarrow P^{\text{H I, g}} \exp[-k^2 \mu^2 (c\sigma_z)^2 / H(z)^2 / 2]$$

$$P^{\text{gg}} \rightarrow P^{\text{gg}} \exp[-k^2 \mu^2 (c\sigma_z)^2 / H(z)^2]$$

in the calculation of V_{eff} . The fact that radial information is lost means that $\beta_{\text{H I}}$ is not going to be measured well, so our primary focus here is measuring the amplitude $\Omega_{\text{H I}} b_{\text{H I}}$.

We will assume a Λ CDM expansion history and keep all cosmological parameters fixed to the Planck 2015 cosmology, in order to examine what constraints can be obtained on a minimal set of parameters. Furthermore, we will suppose that the galaxy bias b_g is known (i.e. measured from the galaxy survey). We therefore parametrize $P^{\text{H I, g}}$ using two parameters, namely ($\Omega_{\text{H I}} b_{\text{H I}}$, $\beta_{\text{H I}}$), and use the aforementioned fiducial models for $\Omega_{\text{H I}}$ and $b_{\text{H I}}$ in our forecasts. We present the marginalized constraints for $\Omega_{\text{H I}} b_{\text{H I}}$ in Table 4. Even if β is poorly measured, we can constrain $\Omega_{\text{H I}} b_{\text{H I}}$ at an ~ 5 per cent level across a wide range of redshift. This is still much better than the currently available constraints on $\Omega_{\text{H I}} b_{\text{H I}}$. We note that while the power spectrum approach in this section is convenient for predicting parameter constraints, in practice an analysis of photometric redshift survey data is likely to use the angular (2D) galaxy clustering power spectrum (Nock, Percival & Ross 2010; Ross et al. 2011). In this way, RSD information can be recovered from cross-correlations between different redshift bins (Asorey, Crocce & Gaztanaga 2014).

It is worth commenting on how the above constraints change if we assume a much smaller, spectroscopic-like redshift error, which we take to be one order of magnitude smaller than the DES-like error, i.e. we consider $\sigma_z = 0.005$. We find that there is a factor of ~ 2 improvement in the $\Omega_{\text{H I}} b_{\text{H I}}$ constraints across the L band, and a smaller improvement at higher redshifts. The biggest effect of course comes in measuring β , as radial information is now retained: we get an ~ 10 per cent fractional error from $z = 0.2$ to 0.8, which corresponds to a measurement of the H I density since the

Table 4. Forecasted fractional uncertainties on $\Omega_{\text{H I}} b_{\text{H I}}$, assuming MeerKAT and a Stage III photometric survey. The sky overlap is taken to be 4000 deg².

z	$\delta(\Omega_{\text{H I}} b_{\text{H I}}) / (\Omega_{\text{H I}} b_{\text{H I}})$
L band	
0.1	0.09
0.2	0.05
0.3	0.04
0.4	0.04
0.5	0.04
UHF band	
0.6	0.05
0.7	0.05
0.8	0.06
0.9	0.07
1.0	0.07
1.1	0.09
1.2	0.10
1.3	0.11
1.4	0.12

Table 5. Forecasted fractional uncertainties on $\{f\sigma_8, D_A, H\}$ at each redshift bin, assuming the SKA1 IM and Stage IV spectroscopic survey specifications described in the main text.

z	$\delta(f\sigma_8) / (f\sigma_8)$	$\delta D_A / D_A$	$\delta H / H$
UHF band			
0.6	–	–	–
0.7	0.04	0.03	0.02
0.8	0.05	0.03	0.02
0.9	0.05	0.03	0.03
1.0	0.06	0.04	0.03
1.1	0.07	0.04	0.03
1.2	0.08	0.05	0.03
1.3	0.10	0.06	0.03
1.4	0.11	0.06	0.04

degeneracy between $\Omega_{\text{H I}}$ and $b_{\text{H I}}$ is broken (as we described in the previous section).

Modifying our approach as before, we now suppose that \bar{T}_b is known and we focus on measuring the growth of structure across redshift. We will consider the combination of a Stage IV spectroscopic survey with an IM survey performed using SKA1, assuming a 7000 deg² overlap and 4000 h total observation time. We write the cross-correlation power spectrum as

$$P^{\text{H I, g}}(k_f, \mu_f, z) = \frac{D_A(z)_f^2 H(z)}{D_A(z)^2 H(z)_f} (b_{\text{H I}} \sigma_8(z) + f \sigma_8(z) \mu^2) \times (b_g \sigma_8(z) + f \sigma_8(z) \mu^2) \frac{P(k, z)}{\sigma_8^2(z)}, \quad (26)$$

and we suppose that $b_g \sigma_8$ has been measured from the galaxy survey, so we keep it fixed. We will therefore use the parameter set $\{\ln f \sigma_8, \ln b_{\text{H I}} \sigma_8, \ln D_A, \ln H\}$ in our Fisher matrix. We summarize the (marginalized) forecasted constraints in Table 5. We note that we have checked our Fisher matrix code results are in agreement with the ones presented in Majerotto et al. (2016) when we consider

a Stage IV galaxy survey only and constrain $\{\ln f\sigma_8, \ln b_g\sigma_8, \ln D_A, \ln H\}$ using the $P^{\text{sg}}(k, \mu, z)$ power spectrum.

We see that we get very good constraints on the growth of structure, the Hubble parameter and the angular diameter distance across a wide range of redshifts corresponding to the era where dark energy is starting to become important ($0.7 < z < 1.4$). Once again, we stress that the above forecasts are expected to be more trustworthy and robust than the autocorrelation ones presented in Table 2, due to the alleviation of systematic effects.

5 PROSPECTS FOR ISW DETECTION WITH INTENSITY MAPPING SURVEYS

In this section, we will consider the possibility of detecting the ISW effect in cross-correlation using neutral hydrogen (H I) as a tracer of the underlying LSS. The late ISW effect (Sachs & Wolfe 1967) arises from the time variation of the gravitational potential and can be detected by cross-correlating the CMB with a low-redshift tracer of the matter distribution (see e.g. Fosalba, Gaztanaga & Castander 2003; Boughn & Crittenden 2005; Giannantonio et al. 2008, 2012). Detection of the ISW effect in a flat Universe provides direct evidence for dark energy.

5.1 The ISW effect in cross-correlation

Let us assume some tracer of matter (tr) – this can be e.g. galaxies (g) or neutral hydrogen (H I). The tracer density contrast we observe in a direction \hat{n}_1 will be

$$\delta_{\text{tr}}(\hat{n}_1) = \int b_{\text{tr}}(z)W(z)\delta_{\text{m}}(\hat{n}_1, z)dz, \quad (27)$$

where $W(z)$ is the selection function of the survey, $b_{\text{tr}}(z)$ is the tracer bias and δ_{m} is the matter density perturbations. The observed tracer density will be correlated with the ISW temperature fluctuation in a direction \hat{n}_2 , which is

$$\frac{\Delta T}{T}(\hat{n}_2) = -2 \int e^{-\tau(z)} \frac{d\Phi}{dz}(\hat{n}_2, z)dz, \quad (28)$$

where Φ is the gravitational potential and $e^{-\tau(z)}$ is the visibility function of the photons.

Having a CMB map and a tracer (galaxy or H I) survey, the autocorrelation and cross-correlation power spectra will be given by

$$\mathbf{C}_\ell^{\text{tr-tr}} = 4\pi \int \frac{dk}{k} \Delta^2(k) \mathcal{I}_\ell^{\text{ISW}}(k) \mathcal{I}_\ell^{\text{tr}}(k), \quad (29)$$

$$\mathbf{C}_\ell^{\text{tr-tr}} = 4\pi \int \frac{dk}{k} \Delta^2(k) [\mathcal{I}_\ell^{\text{tr}}(k)]^2, \quad (30)$$

where $\Delta(k)$ is the scale invariant matter power spectrum $\Delta^2(k) \equiv 4\pi k^3 P(k)/(2\pi)^3$, and the two integrands are, respectively,

$$\mathcal{I}_\ell^{\text{ISW}}(k) = -2 \int e^{-\tau(z)} \frac{d\Phi_k}{dz} j_\ell[kr(z)]dz, \quad (31)$$

$$\mathcal{I}_\ell^{\text{tr}}(k) = \int b_{\text{tr}}(z)W(z)D(z)j_\ell[kr(z)]dz, \quad (32)$$

where Φ_k , $\delta_{\text{m}}(k, z)$ are the Fourier components of the gravitational potential and matter perturbations, and j_ℓ are the spherical Bessel functions.

5.2 Using H I as a tracer

Letting $\text{tr} = \text{H I}$, the density contrast can be written in terms of the 21-cm temperature fluctuations, $\delta_{\text{H I}}(\hat{n}_1) \equiv \delta T_{\text{b}}/\bar{T}_{\text{b}}$ – note that

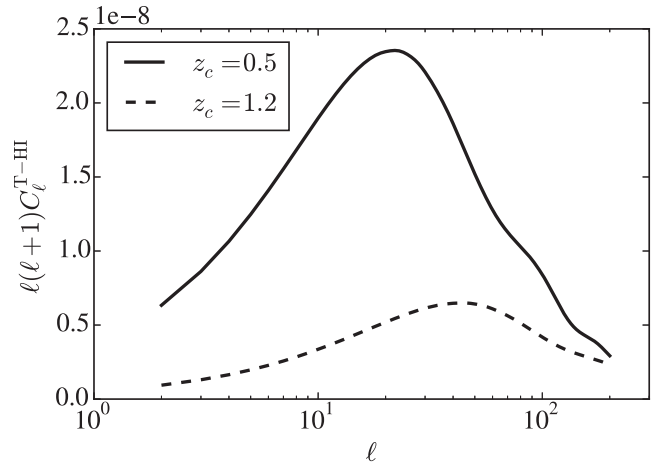


Figure 3. The angular spectrum of the correlation of the *Planck* CMB temperature at $z_c = 0.5$ (solid black line) and $z_c = 1.2$ (dashed black line) for our fiducial cosmology. The bin width is $\Delta z = 0.1$.

we will assume thin enough bins ($\Delta z = 0.1$) so that we can consider \bar{T}_{b} constant within the bin, equal with its value at the central bin redshift z_c . Using the Limber approximation (Limber 1954; LoVerde & Afshordi 2008), we find

$$\mathbf{C}_\ell^{\text{T-H I}} = -\frac{3\Omega_{\text{m},0}(H_0/c)^3}{(\ell+1/2)^2} \int dz W(z)E(z)b_{\text{H I}}(z) \times P_0((\ell+1/2)/r(z))D^2(z)(f(z)-1), \quad (33)$$

$$\mathbf{C}_\ell^{\text{H I-H I}} = \frac{H_0}{c} \int dz E(z)W^2(z)P_0((\ell+1/2)/r(z)) \times b_{\text{H I}}^2(z)D^2(z)/r^2(z), \quad (34)$$

where $E(z) \equiv H(z)/H_0$. Note that if we set $\text{H I} \rightarrow \text{g}$, we recover the literature results for the galaxy case (see e.g. Afshordi 2004; Giannantonio et al. 2008; Francis & Peacock 2010).

The characteristics of the cross-correlation signal have been studied extensively (see Afshordi 2004 for a detailed analysis), but we will reproduce the main results here for completeness. Our fiducial cosmology is the Planck Λ CDM best-fitting model (Planck Collaboration XIII 2016). In Fig. 3, we show the angular power spectrum $\mathbf{C}_\ell^{\text{T-H I}}$ for two redshifts $z_c = 0.5$ and 1.2 taking a $\Delta z = 0.1$ bin in both cases. The ISW cross-correlation is distributed over a wide redshift range $0 < z < 2$, but the main contribution comes from $z \sim 0.4$. Also note that the signal lies on scales larger than a degree, $\ell < 200$.

5.3 Signal-to-noise forecasts

Having the formalism set-up, we can now investigate what can be achieved with forthcoming IM surveys. Using H I as a tracer and an IM survey, the SNR will be given by

$$\left(\frac{S}{N}\right)^2 = f_{\text{sky}} \sum_{\ell=\ell_{\text{min}}}^{\ell_{\text{max}}} (2\ell+1) \frac{[\mathbf{C}_\ell^{\text{T-H I}}]^2}{(\mathbf{C}_\ell^{\text{H I-H I}} + N_{\ell}^{\text{H I}})\mathbf{C}_\ell^{\text{TT}} + [\mathbf{C}_\ell^{\text{T-H I}}]^2}, \quad (35)$$

where $N_{\ell}^{\text{H I}}$ is the instrumental noise of the telescope performing the IM survey, and f_{sky} the (overlapping) fraction of the sky the CMB and IM surveys scan. We have ignored the instrumental noise of the

Table 6. Forecasts of the SNR for detecting the late-time ISW effect using *Planck* and IM surveys.

IM Survey	z range	f_{sky}	SNR
Perfect survey	0–3	1	6.3
BINGO	0.12–0.48	0.05	0.7
MeerKAT UHF band	0.4–1.45	0.1	1.5
SKA1-MID Band 1	0.35–3.06	0.7	4.6
SKA2-like Band 1	0.35–3.06	0.7	4.9

Planck satellite (which would add to $\mathbf{C}_\ell^{\text{TT}}$) because it is negligible, especially on large scales.

Let us first assume that we have the ‘perfect survey’, i.e. set $N_\ell^{\text{H I}} = 0$ and $f_{\text{sky}} = 1$. This will give us the maximum SNR we can achieve. We assume that the survey covers the redshift range $0 < z < 3$ and we will split it in bins with $\Delta z = 0.1$, and take $W(z) = 1/\Delta z$, a top-hat function. For these forecasts, we will set $b_{\text{H I}} = 1$ for simplicity.

The total SNR is calculated by applying equation (35) at each independent redshift bin and then adding the $(S/N)^2$. For our assumed cosmology, we find $(S/N)_{\text{tot}} \sim 6.5$, in agreement with previous results assuming a perfect galaxy survey (Crittenden & Turok 1996; Afshordi 2004). Below we present our results for the various IM surveys we have considered in this paper; the results are also summarized in Table 6. Note that the noise $N_\ell^{\text{H I}}$ in equation (35) is given by (Battye et al. 2013; Pourtsidou et al. 2016)

$$N_\ell^{\text{H I}} = \Omega_{\text{pix}} (T_{\text{sys}}^2 / 2Bt_{\text{obs}}) \exp[\ell(\ell+1)(\theta_B / \sqrt{8\ln 2})^2] / \bar{T}_b^2, \quad (36)$$

with B the frequency bandwidth of observation.

5.3.1 BINGO

Using the formalism outlined above and the BINGO telescope parameters we find $S/N < 1$. This is not surprising, since the BINGO redshift range ($0.12 < z < 0.48$) and sky coverage (2000 deg^2) are not large enough – even if we ignore the instrumental noise $N_\ell^{\text{H I}}$, the SNR is below unity.

5.3.2 MeerKAT

Our best chance for detecting the ISW effect with a 4000 deg^2 MeerKAT survey (in cross-correlation with *Planck*) is performing the IM survey in the UHF band, which corresponds to the redshift range $0.4 < z < 1.45$. We find $S/N = 1.5$ and, again, neglecting the instrumental noise $N_\ell^{\text{H I}}$ only slightly increases the significance of the detection.

5.3.3 SKA-Mid, Phases 1 and 2

In single-dish mode, SKA1-Mid can perform an all-sky IM survey with useful area $f_{\text{sky}} = 0.7$. We are going to consider Band 1 (B1), with $0.35 < z < 3.06$ in our forecasts, and an observation time of 5000 h. We find $S/N = 4.6$. Considering an SKA2-like survey – with an order of magnitude lower thermal (instrumental) noise – we find $S/N = 4.9$, which is practically the maximum SNR that can be achieved by a survey with $f_{\text{sky}} = 0.7$ and redshift range $0.35 < z < 3.06$. This is competitive with the forecasted performance of next-generation LSS optical galaxy surveys like *Euclid*

and LSST⁷ (Douspis et al. 2008; Refregier et al. 2010). Finally, we note that the possibility of detecting the ISW effect using high-redshift ($z \sim 30$) 21-cm maps cross-correlated with galaxies has been studied in Raccanelli et al. (2016), where it was shown that a good detection requires very advanced radio instruments (for example, lunar interferometers).

6 CONCLUSIONS

In this paper, we have shown how ongoing and near-term IM and optical galaxy surveys can be used to constrain the Universe’s expansion history and the growth of cosmic structure, as well as the evolution of H I. We forecasted H I and cosmological constraints for a range of H I surveys and their cross-correlation with galaxies, with special emphasis on the performance of the SKA and its pathfinder MeerKAT.

Our autocorrelation forecasts show that precision measurements of H I and cosmological parameters can be performed already with an $\sim 4000 \text{ deg}^2$ IM survey with MeerKAT, across a wide range of redshift. Constraining the H I evolution and bias tomographically before the SKA comes online is very important for maximizing its scientific output, and probing the H I abundance and evolution across cosmic time is of key importance for astrophysics and cosmology alike. Our cross-correlation forecasts establish that measurements of exquisite precision can be made combining 21-cm intensity maps and optical galaxies, with the extra advantage of alleviating major issues like systematic effects and foreground contaminants, which are relevant for one type of survey but not for the other. Finally, we showed that a large sky survey with Phase 1 of the SKA combined with the *Planck* temperature maps can detect the ISW effect at a level competitive with state-of-the-art Stage IV optical galaxy surveys.

We believe that the results of this paper provide strong motivation for exploring further the possibilities of cross-correlations between 21-cm IM, optical galaxies and the CMB.

ACKNOWLEDGEMENTS

AP’s research is supported by a Dennis Sciama Fellowship at the University of Portsmouth. We acknowledge use of the *CAMB* code (Lewis, Challinor & Lasenby 2000) and *ASTROPY* (Astropy Collaboration 2013). We would like to thank Roy Maartens, Andrej Obuljen and Mario Santos for useful discussions. Fisher matrix codes used in this work are available from <https://github.com/Alkistis/IM-Fish>.

REFERENCES

- Afshordi N., 2004, *Phys. Rev. D*, 70, 083536
Alonso D., Bull P., Ferreira P. G., Santos M. G., 2015, *MNRAS*, 447, 400
Anderson L. et al., 2014, *MNRAS*, 441, 24
Ansari R. et al., 2012, *A&A*, 540, A129
Asorey J., Crocce M., Gaztanaga E., Lewis A., 2012, *MNRAS*, 427, 1891
Asorey J., Crocce M., Gaztanaga E., 2014, *MNRAS*, 445, 2825
Astropy Collaboration, 2013, *A&A*, 558, A33
Aubourg É. et al., 2015, *Phys. Rev. D*, 92, 123516
Battye R. A., Davies R. D., Weller J., 2004, *MNRAS*, 355, 1339
Battye R., Browne I., Dickinson C., Heron G., Maffei B., Pourtsidou A., 2013, *MNRAS*, 434, 1239
Betoule M. et al., 2014, *A&A*, 568, A22

⁷ www.lsst.org

- Bigot-Sazy M.-A. et al., 2016, in Qain L., Li D., eds, ASP Conf. Ser. Vol. 502, Frontiers in Radio Astronomy and FAST Early Sciences Symposium 2015. Astron. Soc. Pac., San Francisco, p. 41
- Blake C., Bridle S., 2005, MNRAS, 363, 1329
- Blake C. et al., 2011, MNRAS, 418, 1707
- Boughn S. P., Crittenden R. G., 2005, New Astron. Rev., 49, 75
- Bull P., Ferreira P. G., Patel P., Santos M. G., 2015, ApJ, 803, 21
- Castorina E., Villaescusa-Navarro F., 2016, preprint ([arXiv:1609.05157](https://arxiv.org/abs/1609.05157))
- Chang T.-C., Pen U.-L., Peterson J. B., McDonald P., 2008, Phys. Rev. Lett., 100, 091303
- Chang T.-C., Pen U.-L., Bandura K., Peterson J. B., 2010, Nature, 466, 463
- Chevallier M., Polarski D., 2001, Int. J. Mod. Phys., D10, 213
- Crichton N. H. M. et al., 2015, MNRAS, 452, 217
- Crittenden R. G., Turok N., 1996, Phys. Rev. Lett., 76, 575
- Delubac T. et al., 2015, A&A, 574, A59
- Douspis M., Castro P. G., Caprini C., Aghanim N., 2008, A&A, 485, 395
- Feldman H. A., Kaiser N., Peacock J. A., 1994, ApJ, 426, 23
- Fosalba P., Gaztanaga E., Castander F., 2003, ApJ, 597, L89
- Francis C. L., Peacock J. A., 2010, MNRAS, 406, 2
- Giannantonio T., Scranton R., Crittenden R. G., Nichol R. C., Boughn S. P., Myers A. D., Richards G. T., 2008, Phys. Rev. D, 77, 123520
- Giannantonio T., Crittenden R., Nichol R., Ross A. J., 2012, MNRAS, 426, 2581
- Khandai N., Sethi S. K., Di Matteo T., Croft R. A. C., Springel V., Jana A., Gardner J. P., 2011, MNRAS, 415, 2580
- Komatsu E. et al., 2011, ApJS, 192, 18
- Kowalski M. et al., 2008, ApJ, 686, 749
- Lampeitl H. et al., 2009, MNRAS, 401, 2331
- Lewis A., Challinor A., Lasenby A., 2000, ApJ, 538, 473
- Limber D. N., 1954, ApJ, 119, 655
- Linder E. V., 2003, Phys. Rev. Lett., 90, 091301
- Linder E. V., Cahn R. N., 2007, Astropart. Phys., 28, 481
- Liu A., Tegmark M., 2011, Phys. Rev. D, 83, 103006
- Loeb A., Wyithe S., 2008, Phys. Rev. Lett., 100, 161301
- LoVerde M., Afshordi N., 2008, Phys. Rev. D, 78, 123506
- McPherson A., 2015, Report and Options for Re-baselining of SKA-1
- Majerotto E., Sapone D., Schäfer B. M., 2016, MNRAS, 456, 109
- Mao Y., Tegmark M., McQuinn M., Zaldarriaga M., Zahn O., 2008, Phys. Rev. D, 78, 023529
- Martin A. M., Papastergis E., Giovanelli R., Haynes M. P., Springob C. M., Stierwalt S., 2010, ApJ, 723, 1359
- Masui K. W., McDonald P., Pen U.-L., 2010, Phys. Rev. D, 81, 103527
- Masui K. et al., 2013, ApJ, 763, L20
- Nock K., Percival W. J., Ross A. J., 2010, MNRAS, 407, 520
- Noterdaeme P. et al., 2012, A&A, 547, L1
- Olivari L. C., Remazeilles M., Dickinson C., 2016, MNRAS, 456, 2749
- Padmanabhan H., Refregier A., 2017, MNRAS, 464, 4008
- Padmanabhan H., Choudhury T. R., Refregier A., 2015, MNRAS, 447, 3745
- Perlmutter S. et al., 1999, ApJ, 517, 565
- Peterson J. B. et al., 2009, preprint ([arXiv:0902.3091](https://arxiv.org/abs/0902.3091))
- Planck Collaboration XIII, 2016, A&A, 594, A13
- Pourtsidou A., Bacon D., Crittenden R., Metcalf R. B., 2016, MNRAS, 459, 863
- Raccanelli A., Kovetz E., Dai L., Kamionkowski M., 2016, Phys. Rev. D, 93, 083512
- Rao S. M., Turnshek D. A., Nestor D., 2006, ApJ, 636, 610
- Rassat A. et al., 2008, preprint ([arXiv:0810.0003](https://arxiv.org/abs/0810.0003))
- Refregier A., Amara A., Kitching T. D., Rassat A., Scaramella R., Weller J., 2010, preprint ([arXiv:1001.0061](https://arxiv.org/abs/1001.0061))
- Rhee J., Zwaan M. A., Briggs F. H., Chengalur J. N., Lah P., Oosterloo T., van der Hulst T., 2013, MNRAS, 435, 2693
- Riess A. G. et al., 1998, AJ, 116, 1009
- Ross A. J., Percival W. J., Crocce M., Cabre A., Gaztanaga E., 2011, MNRAS, 415, 2193
- Sachs R. K., Wolfe A. M., 1967, ApJ, 147, 73
- Samushia L. et al., 2011, MNRAS, 410, 1993
- Santos M. G. et al., 2015, in Proc. Advancing Astrophysics with the Square Kilometre Array, Vol. 019, Cosmology with a SKA HI Intensity Mapping Survey
- Sarkar D., Bharadwaj S., Anathpindika S., 2016, MNRAS, 460, 4310
- Seo H.-J., Eisenstein D. J., 2007, ApJ, 665, 14
- Seo H.-J., Dodelson S., Marriner J., McGinnis D., Stebbins A., Stoughton C., Vallinotto A., 2010, ApJ, 721, 164
- Smith R. et al., 2003, MNRAS, 341, 1311
- Switzer E. et al., 2013, MNRAS, 434, L46
- Switzer E. R., Chang T.-C., Masui K. W., Pen U.-L., Voytek T. C., 2015, ApJ, 815, 51
- Tegmark M., 1997, Phys. Rev. Lett., 79, 3806
- Villaescusa-Navarro F. et al., 2016, MNRAS, 456, 3553
- White M., Song Y.-S., Percival W. J., 2008, MNRAS, 397, 1348
- Wolz L., Abdalla F. B., Blake C., Shaw J. R., Chapman E., Rawlings S., 2014, MNRAS, 441, 3271
- Wolz L. et al., 2017, MNRAS, 464, 4938
- Wolz L., Tonini C., Blake C., Wyithe J. S. B., 2016, MNRAS, 458, 3399
- Zwaan M. A., Meyer M. J., Staveley-Smith L., Webster R. L., 2005, MNRAS, 359, L30

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.