

Implications from clean observables for the binned analysis of $B \rightarrow K^* \mu^+ \mu^-$ at large recoil

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ABSTRACT: We perform a frequentist analysis of q^2 -dependent $B \rightarrow K^*(\rightarrow K\pi)\ell^+\ell^-$ angular observables at large recoil, aiming at bridging the gap between current theoretical analyses and the actual experimental measurements. We focus on the most appropriate set of observables to measure and on the role of the q^2 -binning. We highlight the importance of the observables P_i exhibiting a limited sensitivity to soft form factors for the search for New Physics contributions. We compute predictions for these binned observables in the Standard Model, and we compare them with their experimental determination extracted from recent LHCb data. Analysing $b \rightarrow s$ and $b \rightarrow s\ell^+\ell^-$ transitions within four different New Physics scenarios, we identify several New Physics benchmark points which can be discriminated through the measurement of P_i observables with a fine q^2 -binning. We emphasise the importance (and risks) of using observables with (un)suppressed dependence on soft form factors for the search of New Physics, which we illustrate by the different size of hadronic uncertainties attached to two related observables (P_1 and S_3). We illustrate how the q^2 -dependent angular observables measured in several bins can help to unravel New Physics contributions to $B \rightarrow K^*(\rightarrow K\pi)\ell^+\ell^-$, and show the extraordinary constraining power that the clean observables will have in the near future. We provide semi-numerical expressions for these observables as functions of the relevant Wilson coefficients at the low scale.

KEYWORDS: Rare Decays, B-Physics, Beyond Standard Model

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1 Introduction

The set of rare B -meson decays mediated by the $b \rightarrow s$ transition has been thoroughly studied for many years both from the theoretical and experimental sides. Recently, this set has been complemented by increasingly precise experimental measurements of q^2 -dependent angular observables in the decay $B \rightarrow K^*(\rightarrow K\pi)\ell^+\ell^-$ [1–5]. This has triggered extensive theoretical work studying the constraining power of radiative and semileptonic B decays on New Physics in the framework of the weak effective Hamiltonian [6–14]. These constraints apply mostly to the Wilson coefficients $\mathcal{C}_7^{(\prime)}$, $\mathcal{C}_9^{(\prime)}$ and $\mathcal{C}_{10}^{(\prime)}$ related to the magnetic and semileptonic operators \mathcal{O}_7 , \mathcal{O}_9 , \mathcal{O}_{10} and the corresponding chirality-flipped operators

(which are highly suppressed in the Standard Model). In addition, the tight experimental bounds set in the last few months on the differential decay rate of $B_s \rightarrow \mu^+\mu^-$ [15–18], pushed it close to its small Standard Model (SM) prediction [19, 20] — the theoretical and experimental values are actually brought even closer by the B_s mixing correction $O(\Delta\Gamma_s)$ to B_s branching ratios measured at LHCb, discussed in refs. [21, 22] and applied to $B_s \rightarrow \mu^+\mu^-$ in ref. [23]. This puts strong constraints on $\mathcal{C}_{10}^{(\prime)}$, as well as on the coefficients $\mathcal{C}_S^{(\prime)}$ and $\mathcal{C}_P^{(\prime)}$ of scalar and pseudoscalar operators, specially when combined with $B \rightarrow K\ell^+\ell^-$ data [24, 25].

Small experimental errors and a good control over hadronic uncertainties on the theory side are the key ingredients for these constraints to be efficient. At present, the branching ratio of the inclusive radiative decay $B \rightarrow X_s\gamma$, and the CP asymmetry of $B \rightarrow K^*\gamma$ constitute the strongest constraints on the $\mathcal{C}_7\text{-}\mathcal{C}'_7$ plane. However, the complementarity of constraints among different observables can be exploited to reduce considerably the parameter space. In this respect, the inclusion of the isospin asymmetry of $B \rightarrow K^*\gamma$, together with the forward-backward asymmetry A_{FB} and the longitudinal polarization fraction F_L in $B \rightarrow K^*\ell^+\ell^-$ integrated over the dilepton invariant mass q^2 between 1 and 6 GeV² already impose additional nontrivial constraints on the $\mathcal{C}_7\text{-}\mathcal{C}'_7$ plane, as well as on $\mathcal{C}_9^{(\prime)}$ and $\mathcal{C}_{10}^{(\prime)}$ [6, 7].

This complementarity can be exploited further by considering the q^2 -dependence of angular $B \rightarrow K^*\ell^+\ell^-$ observables. Indeed, increasingly precise measurements of these observables integrated in smaller bins are being presented, due in part to the important statistics obtained from the large data sets collected at the LHCb experiment. This in turn allows for more complete angular analyses providing more observables [4]. The prospects for the near future are very good, aiming towards a complete angular analysis with a fine q^2 -binning.

At this point, it becomes crucial to handle theoretical uncertainties as accurately as possible. The lack of huge deviations in B physics up to now forces us to be precise and conservative in our theoretical predictions. Before claiming any discrepancy, indication or discovery of New Physics, one must be sure that a description in terms of observables with little dependence on the specific choice of hadronic parameters has been used. While the uncertainties related to the form factors constitute a dominant part of the theoretical error, there is a wide spread of quoted uncertainties for $B \rightarrow K^*$ form factors in the recent literature, ranging from a $\sim 10\%$ to a $\sim 40\%$ error for the same form factor [26, 27]. The size of this error does not depend only on the particular theoretical method used to compute the form factor, but also on the delicate estimation of errors associated to the assumptions built in each procedure. For example, the values $A_0(0) = 0.33 \pm 0.03$ and $V(0) = 0.31 \pm 0.04$ given in ref. [26] should be compared to the values $A_0(0) = 0.29 \pm 0.10$ and $V(0) = 0.36 \pm 0.17$ as quoted in ref. [27]. Even central values have shifted significantly, see for instance the value $V(0) = 0.41 \pm 0.05$ from ref. [28] before its update of ref. [26] (also consistent with ref. [29]). Without attempting to discuss the related conceptual issues in any further depth, it is clear that the impact of such discrepancies between different groups concerning hadronic uncertainties can be reduced greatly if one selects quantities that show

a suppressed dependence on the soft form factors, such as the observables P_i or A_T^i defined in refs. [30–32]. These observables can be considered as being theoretically clean in the kinematic range of interest here.

The construction of theoretically clean observables in $B \rightarrow K^*\ell^+\ell^-$ has been the subject of theoretical work for some time [30–35]. Based on the symmetries of the $B \rightarrow K^*\ell^+\ell^-$ angular distribution discussed in ref. [33], a complete characterization of the full distribution in terms of a minimal basis of clean observables has recently been proposed [32]. The conclusion is that, a complete description of the differential decay rate in the limit of massless leptons and in the absence of scalar contributions can be achieved through a set of 6 clean observables $P_{1,2,3,4,5,6}$ complemented by two observables with a significant sensitivity to form factors, e.g., the differential decay rate and the forward-backward asymmetry A_{FB} (or equivalently F_L).¹ In a short term, these clean observables (P_i) should play a leading role in detecting deviations from the SM in a safe way, relegating less clean observables such as $\langle A_{\text{FB}} \rangle$, $\langle F_L \rangle$ or $\langle S_3 \rangle$ to a secondary role of useful cross-checks.

Measurements of the transverse asymmetry $A_T^{(2)}$ have been already provided by the CDF collaboration [2]. However, there has recently been some reluctance from the experimental side to extract such clean observables because their theoretical predictions have been mostly presented as functions of q^2 , while the experimental results are obtained integrated in q^2 -bins. This issue becomes relevant when a coefficient in the angular distribution is expressed as a product of various observables. For example, in ref. [4], the authors prefer to fit for F_L and S_3 instead of F_L and $A_T^{(2)}$ (where $S_3 \sim (1 - F_L)A_T^{(2)}$), arguing that a rapid variation of both F_L and $A_T^{(2)}$ with q^2 could result in a biased estimate of $A_T^{(2)}$ when averaging over large q^2 -bins. This is a perfectly correct statement if one wishes to compare the experimental measurements with (differential) q^2 -dependent observables such as $A_T^{(2)}(q^2)$. Actually, in the specific case of $A_T^{(2)}$ (for the SM case as well as for particular NP values of the Wilson coefficients involved), we observed a very good agreement between the observable integrated naively over the low- q^2 bins (i.e., $\int_{\text{bin}} dq^2 A_T^{(2)}$) and its value derived from the binned observables as they are being measured (denoted $\langle A_T^{(2)} \rangle_{\text{bin}}$ and defined in the next section), showing that the bias is small in this particular case. Though encouraging, this remark will not prevent us from trying to simplify the comparison between theory and experiment by providing theoretical predictions for the *exact* integrated quantities as measured in experiments. These kind of integrated observables have been already discussed in the context of other transverse asymmetries in refs. [11, 35].

In the present paper we address precisely how to analyse efficiently the LHCb measurements on $B \rightarrow K^*\ell^+\ell^-$ at large recoil by choosing a set of clean observables integrated over q^2 -bins. We begin by building a set of integrated observables that correspond in the limit of small binning to the observables in ref. [32], with the aim of making contact with experimentally measurable quantities. In sections 2 and 3, we present Standard Model predictions for these integrated observables, providing the results for different choices of the

¹Mass and scalar effects can be taken into account by including 4 extra clean observables (see ref. [32]). The explicit form of the symmetries in the presence of mass terms and scalar operators can be found in the same paper.

q^2 -binning. In section 4 we perform a model-independent analysis setting constraints on the Wilson coefficients $\mathcal{C}_7^{(\prime)}$, $\mathcal{C}_9^{(\prime)}$ and $\mathcal{C}_{10}^{(\prime)}$, using data from radiative B decays and including the forward-backward asymmetry and the longitudinal polarization fraction of $B \rightarrow K^* \ell^+ \ell^-$, both integrated in the full low- q^2 bin $[1, 6] \text{ GeV}^2$. This updates the analysis of refs. [6, 7] (and related analyses) with several improvements. We use these results to identify a number of New Physics “benchmark points” that are allowed at the 95.5% confidence level by all the constraints considered. In section 5 we discuss the potential complementarity of q^2 -dependent observables in $B \rightarrow K^* \ell^+ \ell^-$ by analyzing our set of clean observables within the NP scenarios specified by the benchmark points. This analysis indicates the scenarios that are more likely to be affected predominantly by the binned observables in $B \rightarrow K^* \ell^+ \ell^-$.

In section 6 we open up a parenthesis to discuss the impact of hadronic uncertainties on the different observables, and the resulting model-independent constraints that follow from different choices of observables. We demonstrate the advantages of using a complete set of clean observables as the one introduced in section 2 and ref. [32].

We then address the model-independent constraints from q^2 -dependent observables in section 7. We extract the experimental values for the three clean observables (P_1, P_2 and P_3) that can be related to the measurements provided in ref. [4]. Our determinations are affected by uncertainties that could be considerably improved, since we lack the experimental information concerning the correlation between the LHCb measurements, which is essential to assess uncertainties in a proper and accurate way. The constraints from q^2 -dependent observables are studied in section 8. First we consider the constraints from the measured A_{FB} and F_L in the two q^2 -bins $[2, 4.3], [4.3, 8.68] \text{ GeV}^2$, then we turn to the constraints imposed by the clean observables $P_{1,2,3}$ in the same bins. In section 9, we compare briefly our study with other similar works.

After concluding in section 10, we include in appendix A the expression of the coefficient J_8 in term of observables and in appendix B a collection of the relevant formulas used to derive the New Physics constraints. In particular we provide the explicit expressions of the clean integrated observables used throughout the paper, for the different choices of q^2 -binning. In appendix C we describe the statistical approach used in the fits.

As a summary of the most important findings, we anticipate the following conclusions of this work:

- Three ($P_{1,2,3}$) out of the six clean observables describing the massless distribution can already be extracted from current measurements, as shown in section 7. While P_3 sets no relevant constraints yet, P_1 and P_2 are complementary to other radiative and semileptonic observables. Combining the measurements of P_2 in two different bins leads to a mild tension with respect to the SM, compatible at 95.5% C.L. but pointing towards a negative contribution to the Wilson coefficient \mathcal{C}_7 .
- The explicit form of the coefficients of the massless angular distribution in terms of the basis of observables is given in eq. (2.13). The expressions turn out to be very simple and exhibit two important features:

- A more natural basis devised to extract information from the distribution in a clean way emerges in the massless case, with a slight redefinition of the observables $P_{4,5,6} \rightarrow P'_{4,5,6}$, which are also clean and defined in eqs. (2.14)–(2.16) (in the SM $P_{4,5,6} \sim P'_{4,5,6}$ to a very good approximation).
- The clean observables in the natural basis ($P_{1,2,3}, P'_{4,5,6}$) can be related to form-factor sensitive observables S_i through the following simple rule:

$$\frac{S_{3,6,9}}{F_T} \rightarrow P_{1,2,3}, \quad \frac{S_{4,5,7}}{\sqrt{F_T F_L}} \rightarrow P'_{4,5,6}, \quad (1.1)$$

where the exact relationships are given in eqs. (2.4) and (2.17).

- The “flipped-sign solution” for C_7 is in general disfavoured by present data at the 95.5% confidence level depending on the NP scenario considered. The isospin asymmetry in $B \rightarrow K^* \gamma$ plays an important role (independent of Wilson coefficients other than $C_7^{(\prime)}$), as well as the forward-backward asymmetry in $B \rightarrow K^* \ell \ell$. This confirms the result of refs. [6, 7].
- We show explicitly the strong impact of the different computations available in the literature for the soft form factors on the theoretical uncertainties for observables like A_{FB}, F_L and S_3 , and the robustness of the clean observables P_i . While the impact on the theoretical error in F_L is evident, the problem for observables like S_3 is more subtle. In the case of S_3 the theoretical uncertainty in the SM is protected by its tiny central value, but away from the SM point the impact can be substantial, preventing this observable from discriminating NP scenarios. None of these problems affect the clean observables P_i or A_T^i .

2 Integrated observables in q^2 -bins

The differential decay rate of the process $\bar{B}_d \rightarrow \bar{K}^*(\rightarrow K\pi)\ell^+\ell^-$ can be written as:

$$\begin{aligned} \frac{d^4\Gamma}{dq^2 d\cos\theta_K d\cos\theta_l d\phi} = & \frac{9}{32\pi} [J_{1s} \sin^2\theta_K + J_{1c} \cos^2\theta_K + (J_{2s} \sin^2\theta_K + J_{2c} \cos^2\theta_K) \cos 2\theta_l \\ & + J_3 \sin^2\theta_K \sin^2\theta_l \cos 2\phi + J_4 \sin 2\theta_K \sin 2\theta_l \cos\phi + J_5 \sin 2\theta_K \sin\theta_l \cos\phi \\ & + (J_{6s} \sin^2\theta_K + J_{6c} \cos^2\theta_K) \cos\theta_l + J_7 \sin 2\theta_K \sin\theta_l \sin\phi + J_8 \sin 2\theta_K \sin 2\theta_l \sin\phi \\ & + J_9 \sin^2\theta_K \sin^2\theta_l \sin 2\phi], \end{aligned} \quad (2.1)$$

where the kinematical variables $\phi, \theta_\ell, \theta_K, q^2$ are defined as in refs. [26, 32, 35]. The decay rate $\bar{\Gamma}$ of the CP-conjugated process $B_d \rightarrow K^*(\rightarrow K\pi)\ell^+\ell^-$ is obtained from eq. (2.1) by replacing $J_{1,2,3,4,7} \rightarrow \bar{J}_{1,2,3,4,7}$ and $J_{5,6,8,9} \rightarrow -\bar{J}_{5,6,8,9}$, where \bar{J} is equal to J with all weak phases conjugated. This corresponds to the same definition of θ_ℓ for both B and \bar{B} (see for example [26, 36]). In this paper we assume that all the observables are CP-averaged, and so are always functions of $J_i + \bar{J}_i$. Therefore, $J_i \rightarrow J_i + \bar{J}_i$ and $\Gamma \rightarrow \Gamma + \bar{\Gamma}$ should be understood in all the formulas below, and in particular all the observables $O(J)$ are assumed to be $O(J + \bar{J})$.

In order to cope with limited statistics, one can write down integrated distributions, such as the uniangular distributions, which depend on a subset of coefficients J_i . This is the way observables such as F_L , A_{FB} or $A_T^{(2)}$ have been measured traditionally. A more recent approach deals with “folded” distributions, with the double advantage of increasing the statistics and focusing on a restricted set of angular coefficients. For example, in ref. [4], the identification of $\phi \leftrightarrow \phi + \pi$ has been used to produce a “folded” angle $\hat{\phi} \in [0, \pi]$ in terms of which a (folded) differential rate $d\hat{\Gamma}(\hat{\phi}) = d\Gamma(\phi) + d\Gamma(\phi - \pi)$ becomes

$$\begin{aligned} \frac{d^4\hat{\Gamma}}{dq^2 d\cos\theta_K d\cos\theta_l d\hat{\phi}} &= \frac{9}{16\pi} \left[J_{1c} \cos^2\theta_K + J_{1s}(1 - \cos^2\theta_K) + J_{2c} \cos^2\theta_K (2 \cos^2\theta_\ell - 1) \right. \\ &\quad + J_{2s}(1 - \cos^2\theta_K)(2 \cos^2\theta_\ell - 1) + J_3(1 - \cos^2\theta_K)(1 - \cos^2\theta_\ell) \cos 2\hat{\phi} \\ &\quad \left. + J_{6s}(1 - \cos^2\theta_K) \cos\theta_\ell + J_9(1 - \cos^2\theta_K)(1 - \cos^2\theta_\ell) \sin 2\hat{\phi} \right]. \end{aligned} \quad (2.2)$$

In the following we will neglect scalar and lepton mass effects. A detailed analysis of the impact of neglecting lepton masses can be found in ref. [37]. Concerning scalar contributions, two observables called S_1 and S_2 were designed in ref. [32] to explore the measurable impact of scalar effects. However, the strong constraint from the $B_s \rightarrow \mu^+\mu^-$ branching ratio already makes these effects negligible. Still it will be interesting to include these corrections once enough statistics is collected.

In this approximation, this distribution can be written as a function of the observables in ref. [32] as follows:

$$\begin{aligned} \frac{d^4\Gamma}{dq^2 d\cos\theta_K d\cos\theta_l d\hat{\phi}} &= \frac{9}{16\pi} \left[F_L \cos^2\theta_K + \frac{3}{4}F_T(1 - \cos^2\theta_K) - F_L \cos^2\theta_K (2 \cos^2\theta_\ell - 1) \right. \\ &\quad + \frac{1}{4}F_T(1 - \cos^2\theta_K)(2 \cos^2\theta_\ell - 1) + \frac{1}{2}P_1F_T(1 - \cos^2\theta_K)(1 - \cos^2\theta_\ell) \cos 2\hat{\phi} \\ &\quad \left. + 2P_2F_T(1 - \cos^2\theta_K) \cos\theta_\ell - P_3F_T(1 - \cos^2\theta_K)(1 - \cos^2\theta_\ell) \sin 2\hat{\phi} \right] \frac{d\Gamma}{dq^2}, \end{aligned} \quad (2.3)$$

where P_1 , P_2 and P_3 are theoretically clean observables [32] that in terms of form factor dependent observables² are given as

$$P_1F_T = 2S_3, \quad P_2F_T = S_{6s}/2, \quad P_3F_T = -S_9, \quad (2.4)$$

or alternatively,³

$$P_2F_T = -2A_{\text{FB}}/3, \quad P_3F_T = -A_{\text{im}}. \quad (2.5)$$

The quantity F_T is defined as $F_T \equiv 1 - F_L$. The observable P_1 is better known by its original name, $A_T^{(2)}$ [30].

²The observables S_i are defined as $S_i = (J_i + \bar{J}_i)/(\bar{\Gamma} + \Gamma)$ [26], while $A_{\text{im}} = S_9$ [31].

³Note that the $\cos\theta_\ell$ term in eq. (2.3) has opposite sign with respect to ref. [4] because of the different definition of the angle θ_ℓ : $\theta_\ell^{\text{as}} = \pi - \theta_\ell^{\text{HCb}}$ for the B decay.

Experimentally, one can fit separately each of the following five independent coefficients that appear in the folded distribution in eq. (2.3):

$$\begin{aligned} c_0(q^2) &= \frac{d\Gamma}{dq^2}, & c_1(q^2) &= P_1 F_T \frac{d\Gamma}{dq^2}, & c_2(q^2) &= P_2 F_T \frac{d\Gamma}{dq^2}, \\ c_3(q^2) &= P_3 F_T \frac{d\Gamma}{dq^2}, & c_4(q^2) &= F_T \frac{d\Gamma}{dq^2}. \end{aligned} \tag{2.6}$$

For each q^2 one can then in principle extract the theoretically clean observables $P_1 \equiv A_T^{(2)}$, P_2 and P_3 , as well as the transverse polarization fraction and the differential decay rate. However, in practice the q^2 -dependence is discretised in a number of bins, and the coefficients $c_i(q^2)$ thus extracted are quantities integrated over particular q^2 -intervals:

$$\langle c_i \rangle_{\text{bin}} = \frac{\int_{\text{bin}} dq^2 c_i(q^2)}{\int_{\text{bin}} dq^2}. \tag{2.7}$$

One could hope to be able to extract also corresponding integrated theoretically clean observables such as

$$\langle A_T^{(2)} \rangle_{\text{bin}}^{(\text{naive})} = \frac{\int_{\text{bin}} dq^2 A_T^{(2)}(q^2)}{\int_{\text{bin}} dq^2}, \tag{2.8}$$

but due to the experimental procedure used, such a determination is achievable asymptotically only, as the bin size goes to zero. The actual theoretically clean quantities that can be extracted from experiment and on which we will focus from now on, must be composed of the integrated quantities in eq. (2.7):

$$\langle P_1 \rangle_{\text{bin}} \equiv \langle A_T^{(2)} \rangle_{\text{bin}} = \frac{\int_{\text{bin}} dq^2 c_1(q^2)}{\int_{\text{bin}} dq^2 c_4(q^2)} = \frac{\langle c_1 \rangle_{\text{bin}}}{\langle c_4 \rangle_{\text{bin}}}, \tag{2.9}$$

$$\langle P_2 \rangle_{\text{bin}} = \frac{\int_{\text{bin}} dq^2 c_2(q^2)}{\int_{\text{bin}} dq^2 c_4(q^2)} = \frac{\langle c_2 \rangle_{\text{bin}}}{\langle c_4 \rangle_{\text{bin}}}, \tag{2.10}$$

$$\langle P_3 \rangle_{\text{bin}} = \frac{\int_{\text{bin}} dq^2 c_3(q^2)}{\int_{\text{bin}} dq^2 c_4(q^2)} = \frac{\langle c_3 \rangle_{\text{bin}}}{\langle c_4 \rangle_{\text{bin}}}. \tag{2.11}$$

Other observables are accessible to the current LHCb data set by means of similar partial angular analyses. Three observables related to P_4 , P_5 and P_6 (see ref. [32]) could be extracted by the LHCb collaboration in the near future. This means that, without actually performing a full angular analysis, the LHCb collaboration could be able to provide measurements of the complete set of 8 observables that describe the full distribution in the massless approximation (six of them being theoretically clean [32]). Therefore, we will work here under the assumption that integrated versions (exactly as in eq. (2.7)) of the observables $c_{0,1,2,3,4}$ and $J_{4,5,7}$ are available experimentally. Any measurable observable must be a combination of the $\langle c_i \rangle_{\text{bin}}$ in eq. (2.7), and of the observables:

$$\langle J_{4,5,7} \rangle_{\text{bin}} = \frac{\int_{\text{bin}} dq^2 J_{4,5,7}(q^2)}{\int_{\text{bin}} dq^2}. \tag{2.12}$$

The coefficients of the angular distribution can be written in terms of the basis of observables (see ref. [32]), and in terms of the coefficients c_i of eq. (2.6) as follows:⁴

$$\begin{aligned}
 J_{1s} &= \frac{3}{4}F_T \frac{d\Gamma}{dq^2} = \frac{3}{4}c_4, & J_{2s} &= \frac{1}{4}F_T \frac{d\Gamma}{dq^2} = \frac{1}{4}c_4 \\
 J_{1c} &= F_L \frac{d\Gamma}{dq^2} = c_0 - c_4, & J_{2c} &= -F_L \frac{d\Gamma}{dq^2} = c_4 - c_0, \\
 J_3 &= \frac{1}{2}P_1 F_T \frac{d\Gamma}{dq^2} = \frac{1}{2}c_1, & J_{6s} &= 2P_2 F_T \frac{d\Gamma}{dq^2} = 2c_2, \\
 J_4 &= \frac{1}{2}P'_4 \sqrt{F_T F_L} \frac{d\Gamma}{dq^2} = \frac{1}{2}P'_4 \sqrt{c_4(c_0 - c_4)}, & J_9 &= -P_3 F_T \frac{d\Gamma}{dq^2} = -c_3, \\
 J_5 &= P'_5 \sqrt{F_T F_L} \frac{d\Gamma}{dq^2} = P'_5 \sqrt{c_4(c_0 - c_4)}, \\
 J_7 &= -P'_6 \sqrt{F_T F_L} \frac{d\Gamma}{dq^2} = -P'_6 \sqrt{c_4(c_0 - c_4)},
 \end{aligned} \tag{2.13}$$

where the primed observables are defined as:

$$P'_4 \equiv P_4 \sqrt{1 - P_1} = \frac{J_4}{\sqrt{-J_{2c} J_{2s}}} \tag{2.14}$$

$$P'_5 \equiv P_5 \sqrt{1 + P_1} = \frac{J_5}{2\sqrt{-J_{2c} J_{2s}}} \tag{2.15}$$

$$P'_6 \equiv P_6 \sqrt{1 - P_1} = -\frac{J_7}{2\sqrt{-J_{2c} J_{2s}}} \tag{2.16}$$

The case of the coefficient J_8 is discussed separately in detail in appendix A. These observables $P'_{4,5,6}$ are clean and coincide to a good approximation with $P_{4,5,6}$ in the SM (due to the fact that $P_1 \simeq 0$ in the SM). The whole analysis can be performed directly in terms of the observables $P_{4,5,6}$; however, from the experimental point of view, fitting the primed observables is simpler and more efficient.

These observables can be related to the observables $S_{4,5,7}$ of ref. [26]:

$$P'_4 = 2 \frac{S_4}{\sqrt{F_T F_L}}, \quad P'_5 = \frac{S_5}{\sqrt{F_T F_L}}, \quad P'_6 = -\frac{S_7}{\sqrt{F_T F_L}}. \tag{2.17}$$

There is therefore no particular advantage for the experimental extraction of the observables S_i instead of the P'_i , while from the theory point of view the P'_i are under better control and suffer from smaller uncertainties.

We want now to construct the theoretically clean integrated observables that correspond to those in ref. [32] (or variations thereof). For $P_{1,2,3}$ the answer is precisely $\langle P_1 \rangle_{\text{bin}}$, $\langle P_2 \rangle_{\text{bin}}$ and $\langle P_3 \rangle_{\text{bin}}$ defined in eqs. (2.9)–(2.11). In analogy with eqs. (2.9)–(2.11), integrated

⁴A generalization of this parameterization including scalars and lepton masses can be found in ref. [32]. An alternative parametrization including lepton masses is given in ref. [37].

versions of the observables $P'_{4,5,6}$ can be defined:

$$\langle P'_4 \rangle_{\text{bin}} = \frac{2 \int_{\text{bin}} dq^2 J_4(q^2)}{\sqrt{\int_{\text{bin}} dq^2 c_4(q^2) \int_{\text{bin}} dq^2 (c_0(q^2) - c_4(q^2))}} = \frac{2 \langle J_4 \rangle_{\text{bin}}}{\sqrt{\langle c_4 \rangle_{\text{bin}} (\langle c_0 \rangle_{\text{bin}} - \langle c_4 \rangle_{\text{bin}})}}, \quad (2.18)$$

$$\langle P'_5 \rangle_{\text{bin}} = \frac{\int_{\text{bin}} dq^2 J_5(q^2)}{\sqrt{\int_{\text{bin}} dq^2 c_4(q^2) \int_{\text{bin}} dq^2 (c_0(q^2) - c_4(q^2))}} = \frac{\langle J_5 \rangle_{\text{bin}}}{\sqrt{\langle c_4 \rangle_{\text{bin}} (\langle c_0 \rangle_{\text{bin}} - \langle c_4 \rangle_{\text{bin}})}}, \quad (2.19)$$

$$\langle P'_6 \rangle_{\text{bin}} = \frac{-\int_{\text{bin}} dq^2 J_7(q^2)}{\sqrt{\int_{\text{bin}} dq^2 c_4(q^2) \int_{\text{bin}} dq^2 (c_0(q^2) - c_4(q^2))}} = \frac{-\langle J_7 \rangle_{\text{bin}}}{\sqrt{\langle c_4 \rangle_{\text{bin}} (\langle c_0 \rangle_{\text{bin}} - \langle c_4 \rangle_{\text{bin}})}}. \quad (2.20)$$

Finally, integrated versions of the longitudinal polarization fraction F_L and the forward-backward asymmetry A_{FB} can be defined in terms of the coefficients c_i in the following way:

$$\langle A_{\text{FB}} \rangle_{\text{bin}} = -\frac{3 \int_{\text{bin}} dq^2 c_2(q^2)}{2 \int_{\text{bin}} dq^2 c_0(q^2)} = -\frac{3 \langle c_2 \rangle_{\text{bin}}}{2 \langle c_0 \rangle_{\text{bin}}}, \quad (2.21)$$

$$\langle F_L \rangle_{\text{bin}} = \frac{\int_{\text{bin}} dq^2 (c_0(q^2) - c_4(q^2))}{\int_{\text{bin}} dq^2 c_0(q^2)} = \frac{\langle c_0 \rangle_{\text{bin}} - \langle c_4 \rangle_{\text{bin}}}{\langle c_0 \rangle_{\text{bin}}}. \quad (2.22)$$

In the following sections we will study these integrated observables in detail.

3 SM predictions for integrated observables

We can provide SM predictions for the set of integrated observables $\langle P_i \rangle$ as well as $\langle A_{\text{FB}} \rangle$ and $\langle F_L \rangle$. In tables 1 and 2 we show the predictions in the q^2 -bins [1,2], [2,4.3], and [4.3,6] (GeV²) — following the binning used by the experimental collaborations up to now (except for the first bin) — as well as the predictions for the integrated low- q^2 observables, in the region [1,6] GeV². The first error accounts for all parametric uncertainties, while the second error corresponds to an estimate of Λ/m_b corrections, as described below. In figures 1, 2 and 3 we show the corresponding SM predictions for the observables in the case of one and three bins (corresponding to the predictions in tables 1 and 2), as well as for five bins with a width of 1 GeV².

The SM predictions are obtained as follows. The observables integrated over each bin are defined in terms of the coefficients $c_i(q^2)$ in eqs. (2.9)–(2.11), (2.18)–(2.20), (2.21) and (2.22). The coefficients $c_i(q^2)$ are simple functions of transversity amplitudes (see for example ref. [32]). The transversity amplitudes can be written in terms of Wilson coefficients and $B \rightarrow K^*$ form factors following refs. [38, 39]. Concerning the Wilson coefficients, the form factors, and the treatment of uncertainties, we proceed as in refs. [6, 7, 32], with a slight revision in the treatment of form factors:

Wilson coefficients. The SM Wilson coefficients are evaluated at the matching scale $\mu_0 = 2M_W$, and evolved down to the hadronic scale $\mu_b = 4.8 \text{ GeV}$ following refs. [40–44]. The running of quark masses and couplings proceeds analogously. The SM Wilson coefficients at the scale μ_b are shown in table 3.

q^2 (GeV ²)	[1, 2]	[2, 4.3]	[4.3, 6]	[1, 6]
$\langle P_1 \rangle$	$0.008^{+0.009+0.051}_{-0.005-0.053}$	$-0.051^{+0.010+0.048}_{-0.009-0.050}$	$-0.100^{+0.001+0.049}_{-0.001-0.053}$	$-0.055^{+0.009+0.049}_{-0.008-0.052}$
$\langle P_2 \rangle$	$0.395^{+0.020+0.011}_{-0.021-0.012}$	$0.227^{+0.055+0.014}_{-0.083-0.016}$	$-0.254^{+0.063+0.034}_{-0.068-0.035}$	$0.080^{+0.054+0.020}_{-0.073-0.021}$
$\langle P_3 \rangle$	$-0.003^{+0.001+0.025}_{-0.002-0.028}$	$-0.004^{+0.001+0.023}_{-0.003-0.025}$	$-0.002^{+0.001+0.022}_{-0.002-0.024}$	$-0.003^{+0.001+0.023}_{-0.002-0.025}$
$\langle P'_4 \rangle$	$-0.160^{+0.036+0.024}_{-0.027-0.025}$	$0.570^{+0.067+0.000}_{-0.054-0.002}$	$0.944^{+0.025+0.000}_{-0.025-0.004}$	$0.553^{+0.060+0.004}_{-0.050-0.008}$
$\langle P'_5 \rangle$	$0.369^{+0.044+0.000}_{-0.061-0.002}$	$-0.343^{+0.089+0.043}_{-0.108-0.046}$	$-0.774^{+0.061+0.087}_{-0.059-0.093}$	$-0.353^{+0.081+0.050}_{-0.095-0.053}$
$\langle P'_6 \rangle$	$-0.095^{+0.025+0.012}_{-0.042-0.011}$	$-0.092^{+0.029+0.026}_{-0.045-0.024}$	$-0.074^{+0.027+0.051}_{-0.038-0.046}$	$-0.085^{+0.027+0.033}_{-0.041-0.029}$

Table 1. SM predictions for the clean observables $\langle P_i \rangle$.

q^2 (GeV ²)	[1, 2]	[2, 4.3]	[4.3, 6]	[1, 6]
$\langle A_{\text{FB}} \rangle$	$-0.214^{+0.111+0.003}_{-0.144-0.002}$	$-0.079^{+0.053+0.004}_{-0.065-0.003}$	$0.112^{+0.086+0.017}_{-0.065-0.016}$	$-0.034^{+0.035+0.009}_{-0.033-0.008}$
$\langle F_L \rangle$	$0.638^{+0.185+0.007}_{-0.236-0.006}$	$0.769^{+0.129+0.006}_{-0.194-0.006}$	$0.706^{+0.151+0.004}_{-0.201-0.004}$	$0.719^{+0.149+0.006}_{-0.208-0.006}$

Table 2. SM predictions for $\langle A_{\text{FB}} \rangle$ and $\langle F_L \rangle$.

Form factors. There are seven $B \rightarrow K^*$ form factors: $V(q^2)$, $A_{0,1,2}(q^2)$ and $T_{1,2,3}(q^2)$. Their determination involves the computation at $q^2 = 0$ and the parameterization of the q^2 -dependence. At $q^2 = 0$, these form factors can be obtained from light-cone sum rules with B-meson distribution amplitudes (see ref. [27]). Concerning their dependence on q^2 , ref. [27] provides a conservative and convenient parameterization (the prospects for these form factors from lattice QCD have been discussed in detail in ref. [34]). The soft form factors $\xi_{\parallel,\perp}$ are defined in terms of the full form factors following ref. [39]. The soft form factor $\xi_{\perp}(0)$ at $q^2 = 0$ is obtained directly from $V(0)$ as given in ref. [27]. The form factor $\xi_{\parallel}(0)$ is defined as a combination of the form factors $A_1(0)$ and $A_2(0)$, and in the large-recoil limit and at leading order in α_s it is proportional to $A_0(0)$. We use $A_0(0)$ to fix $\xi_{\parallel}(q^2)$ at $q^2 = 0$ to a good accuracy, and set its q^2 -dependence to reproduce its exact expression in terms of $A_1(q^2)$ and $A_2(q^2)$ using the parameterization of ref. [27]. The numerical inputs used are collected in table 3.

In the present paper we take the form factors of ref. [27] for two reasons. The first one is to be consistent with the analysis of ref. [32]. The second is to be conservative in the treatment of hadronic uncertainties, showing at the same time that clean observables are mostly insensitive to this choice. The use of the value for $V(0)$ from ref. [26] would shift the central values of A_{FB} and F_L in the whole low- q^2 region, while only a mild effect around $q^2 \simeq 6 \text{ GeV}^2$ would be seen in some of the P_i observables. These form factors have much larger uncertainties than those of refs. [26, 28], and translate into large error bars in A_{FB} , F_L and other form factor dependent observables.

This difference in the size of the uncertainties can be partly explained by the approaches taken to apply light-cone sum rules in refs. [26, 28] and [27]. In refs. [26, 28], the sum rules are written using the light-meson distribution amplitudes up to twist 4 and including $\mathcal{O}(\alpha_s)$ corrections. In ref. [27], the sum rules are written for the B -meson distribution

$\mu_b = 4.8 \text{ GeV}$		$\mu_0 = 2M_W$	[45]
$m_B = 5.27950 \text{ GeV}$	[46]	$m_{K^*} = 0.89594 \text{ GeV}$	[46]
$m_{B_s} = 5.3663 \text{ GeV}$	[46]	$m_\mu = 0.105658367 \text{ GeV}$	[46]
$\sin^2 \theta_W = 0.2313$	[46]		
$M_W = 80.399 \pm 0.023 \text{ GeV}$	[46]	$M_Z = 91.1876 \text{ GeV}$	[46]
$\alpha_{\text{em}}(M_Z) = 1/128.940$	[45]	$\alpha_s(M_Z) = 0.1184 \pm 0.0007$	[46]
$m_t^{\text{pole}} = 173.3 \pm 1.1 \text{ GeV}$	[47]	$m_b^{1S} = 4.68 \pm 0.03 \text{ GeV}$	[48]
$m_c^{\overline{\text{MS}}}(m_c) = 1.27 \pm 0.09 \text{ GeV}$	[46]	$m_s^{\overline{\text{MS}}}(2 \text{ GeV}) = 0.101 \pm 0.029 \text{ GeV}$	[46]
$\lambda_{\text{CKM}} = 0.22543 \pm 0.0008$	[49]	$A_{\text{CKM}} = 0.805 \pm 0.020$	[49]
$\bar{\rho} = 0.144 \pm 0.025$	[49]	$\bar{\eta} = 0.342 \pm 0.016$	[49]
$\mathcal{B}(B \rightarrow X_c e \bar{\nu}) = 0.1061 \pm 0.00017$	[45]	$C = 0.58 \pm 0.016$	[45]
$\lambda_2 = 0.12 \text{ GeV}^2$	[45]		
$\Lambda_h = 0.5 \text{ GeV}$	[50]	$f_B = 0.190 \pm 0.004 \text{ GeV}$	[51]
$f_{K^*,\parallel} = 0.220 \pm 0.005 \text{ GeV}$	[26]	$f_{K^*,\perp}(2 \text{ GeV}) = 0.163(8) \text{ GeV}$	[26]
$V(0) = 0.36_{-0.12}^{+0.23}$	[27]	$A_0(0) = 0.29_{-0.07}^{+0.10}$	[27]
$a_{1,\parallel,\perp}(2 \text{ GeV}) = 0.03 \pm 0.03$	[26]	$a_{2,\parallel,\perp}(2 \text{ GeV}) = 0.08 \pm 0.06$	[26]
$\lambda_B(\mu_h) = 0.51 \pm 0.12 \text{ GeV}$	[26]		
$f_{B_s} = 0.227 \pm 0.004 \text{ GeV}$	[51]	$\tau_{B_s} = 1.497 \pm 0.015 \text{ ps}$	[46]

$\mathcal{C}_1(\mu_b)$	$\mathcal{C}_2(\mu_b)$	$\mathcal{C}_3(\mu_b)$	$\mathcal{C}_4(\mu_b)$	$\mathcal{C}_5(\mu_b)$	$\mathcal{C}_6(\mu_b)$	$\mathcal{C}_7^{\text{eff}}(\mu_b)$	$\mathcal{C}_8^{\text{eff}}(\mu_b)$	$\mathcal{C}_9(\mu_b)$	$\mathcal{C}_{10}(\mu_b)$
-0.2632	1.0111	-0.0055	-0.0806	0.0004	0.0009	-0.2923	-0.1663	4.0749	-4.3085

Table 3. Input parameters used in the analysis and Wilson coefficients at μ_b .

amplitudes up to twist 3 — they include the (significant) soft-gluon emission from charm loops not considered in refs. [26, 28], but neglect the radiative corrections included in these references, as well as $1/m_b$ HQET corrections. Therefore the two analyses are only partially comparable, which explains why the quoted uncertainties differ in size. But one should also emphasise that beyond the approach taken, the hadronic inputs, i.e., the models used for the distribution amplitudes, play a crucial role concerning the uncertainties quoted for the form factors: in the case of light mesons [26, 28], the shapes are constrained by results coming from other light-cone sum rules, whereas in the case of the B meson [27], a large range of variation for the shape models is allowed. At any rate, the discrepancy between refs. [26, 28] and ref. [27] is a clear indication that the theoretical uncertainties attached to these observables (i.e., A_{FB} , F_L) in the literature should be considered with a healthy dose of skepticism, and are strongly dependent on the choice of the $B \rightarrow K^*$ form factors. Fortunately, it will be seen in section 6 that the error bars for the clean observables P_i are not affected by this variation of the form factor uncertainties and remain under good theoretical control.

Uncertainties. We consider five main sources of uncertainties: the renormalization scale μ_b , the quark mass ratio $\hat{m}_c \equiv m_c/m_b$, the form factors, the factor that determines the

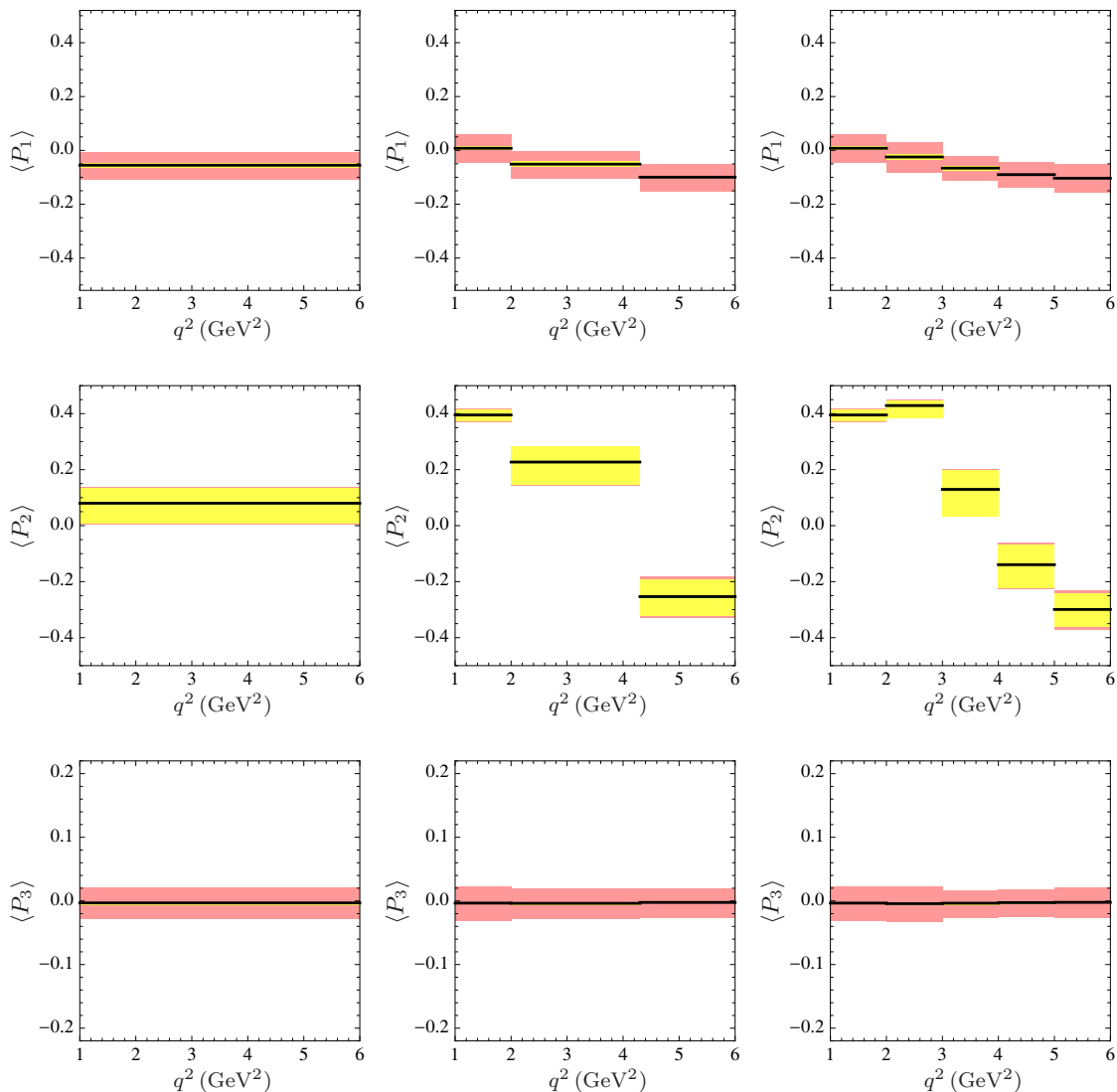


Figure 1. Binned Standard Model predictions for the clean observables $\langle P_{1,2,3} \rangle$, for a single bin $[1, 6] \text{ GeV}^2$ (left column), three bins $[1, 2]$, $[2, 4.3]$, $[4.3, 6] \text{ GeV}^2$ (central column), and five bins of width 1 GeV^2 each (right column). The red (dark gray) error bar correspond to the Λ/m_b corrections, the yellow one (light gray) to the other sources of uncertainties. If one of the two bands is missing, it means the associated uncertainty is negligible compared to the dominant one.

relative size of the hard-scattering term with respect to the form factor contribution to the amplitude (defined in eq. (55) of ref. [38]) and the Λ/m_b power corrections.

We follow the usual procedure consisting in varying the renormalization scale from $\mu_b/2$ to $2\mu_b$. For \hat{m}_c we take $\hat{m}_c = 0.29 \pm 0.02$ (see refs. [38, 52]). Concerning the form factors: we express all the observables as a function of $A_0(q^2)$ and $V(q^2)$. The q^2 -parameterization of these form factors (that of ref. [27]) depends on the normalisation parameters $A_0(0)$, $V(0)$ as well as b_1^V and $b_1^{A_0}$ (encoding the q^2 -dependence of the form factor in the so-called z -parametrisation). We vary simultaneously $V(0)$ and b_1^V , and independently, $A_0(0)$ and $b_1^{A_0}$,

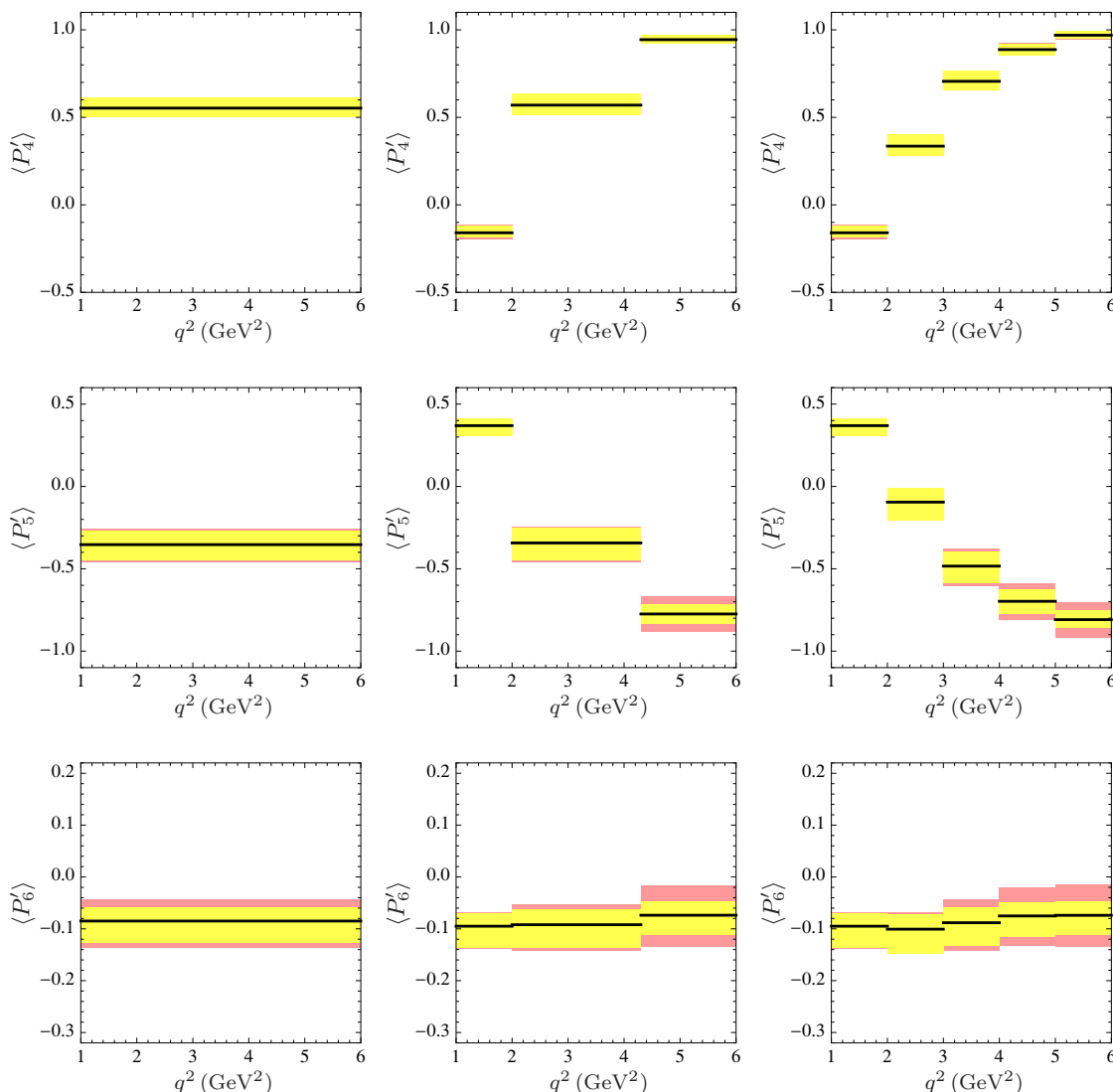


Figure 2. Binned Standard Model predictions for the clean observables $\langle P'_{4,5,6} \rangle$, with the same conventions as in figure 1.

within the errors quoted in table 3 and ref. [27]. The two errors are added in quadrature. Concerning the parameter describing the relative weight of the hard-scattering contribution compared to the form-factor one in ref. [38], its error is estimated at the level of a 25%, where its reduction with respect to ref. [32] (where it was 30%) is due to the updated value of f_B [51]. Finally, for Λ/m_b corrections, we follow the statistical procedure outlined in ref. [33] to produce upper and lower 1σ ranges consistent with a generic 10% contribution of power corrections to the amplitudes. All individual uncertainties are considered separately and their impact on each observable is monitored to produce asymmetric upper and lower errors. All upper and lower uncertainties, excluding Λ/m_b corrections, are added separately in quadrature to produce the first asymmetric error bars shown in tables 1 and 2, and the yellow bands in figures 1, 2 and 3. The second error bars in tables 1 and 2 correspond to

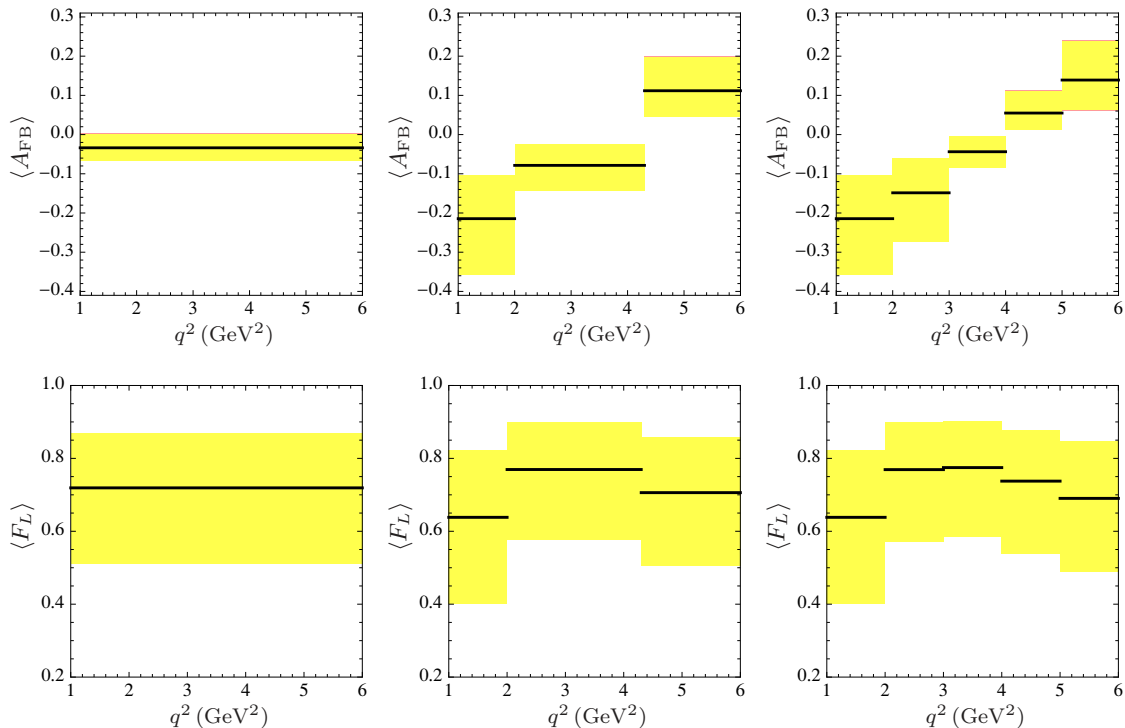


Figure 3. Binned Standard Model predictions for the observables $\langle A_{FB} \rangle$ and $\langle F_L \rangle$, with the same conventions as in figure 1.

the Λ/m_b corrections. In figures 1, 2 and 3, both uncertainties are added linearly to give the larger red error bands.

As can be seen from the plots, some observables appear to be almost insensitive to power corrections. This merely indicates that the hadronic uncertainties are much more important than power corrections for these observables.

While the observables P_i are almost insensitive to the choice of form factors, the uncertainties of other observables vary substantially if the form factors of refs. [26, 28] are used instead. For example, in table 2 we quote the following prediction for F_L : $\langle F_L \rangle_{[1,6]} = 0.719^{+0.149+0.006}_{-0.208-0.006}$, while if we take the form factors of ref. [26] we obtain $\langle F_L \rangle_{[1,6]} = 0.809^{+0.045+0.008}_{-0.032-0.002}$ (see figure 1 of ref. [53]).

It could be wrongly concluded that F_L has a small error in comparison with the predictions for the P_i , if one compares the percentage over the central value of the observable. However, the percentage with respect to the central value is not a sensible measure to compare the size of the errors. In the case of F_L , defined in the range [0,1], the size of the error should be compared with 1. Otherwise, taking instead the observable $F_T = 1 - F_L$ (which has a central value 5 times smaller) one would judge the error as 5 times larger, while the two observables are effectively the same. For the observables P_i , defined in the range [-1, 1] (this is also true approximately for $P'_{4,5,6}$), the same argument applies, and therefore the error percentage should be evaluated over 2 and not the central value. In practice one is concerned with the discriminating power between SM and NP points, which

Observable	Experiment	SM prediction
$\text{BR}(B \rightarrow X_s \gamma)$	$(3.55 \pm 0.26) \cdot 10^{-4}$ [55]	$(3.15 \pm 0.23) \cdot 10^{-4}$ [45]
$S_{K^* \gamma}$	-0.16 ± 0.22 [55]	-0.03 ± 0.01 [6]
$A_I(B \rightarrow K^* \gamma)$	0.052 ± 0.026 [55]	0.041 ± 0.025 [6]
$\text{BR}(B \rightarrow X_s \mu^+ \mu^-)_{[1,6]}$	$(1.60 \pm 0.50) \cdot 10^{-6}$ [56]	$(1.59 \pm 0.11) \cdot 10^{-6}$ [56]
$\langle A_{\text{FB}} \rangle_{[1,6]}$	$-0.13^{+0.068}_{-0.078}$ *	-0.034 ± 0.035 †
$\langle F_L \rangle_{[1,6]}$	$0.622^{+0.059}_{-0.057}$ *	0.719 ± 0.179 †
$\text{BR}(B_s \rightarrow \mu^+ \mu^-)$	$< 4.5 \cdot 10^{-9}$ (at 95.5% C.L.) [16]	$(3.32 \pm 0.17) \cdot 10^{-9}$ [12]

Table 4. Experimental numbers and Standard Model predictions of the observables used in the analysis of section 4. * indicates our own average of the data. † indicates our SM prediction. (See also ref. [19] concerning the $B_s \rightarrow \mu^+ \mu^-$ branching ratio.)

gives further support to this argument: since one expects general (unconstrained) NP to give contributions in the whole range of the observable, the size of the error (as a measure of uncertainty of the position of the SM point) should be compared to the full range.

If this prescription is adopted, the relative errors of the P_i are 2.5%, 3.5%, 1.2%, 3%, 5%, 2.5% for $i = 1 \dots 6$, while for F_L is 18%. These are the results obtained using the form factors in ref. [27]. If instead ref. [26] is used for the form factors, the errors of the P_i barely change, while the relative error for F_L goes down to 5%. This can be seen in figure 1 of ref. [53]. This is one of the benefits of the clean observables: there is no need to rely on complicated estimations of errors in the light-cone sum-rules procedure, as explained above.

The sensitivity to form factor uncertainties of observables such as S_3 is more subtle and will be discussed in section 6.

4 Model-independent constraints without q^2 -binned observables

In this section we revisit the model-independent constraints on the Wilson coefficients \mathcal{C}_7 , \mathcal{C}'_7 , $\mathcal{C}_9^{(\prime)}$ and $\mathcal{C}_{10}^{(\prime)}$ from well controlled observables, excluding all $B \rightarrow K^* \ell^+ \ell^-$ observables except for $\langle A_{\text{FB}} \rangle_{[1,6]}$ and $\langle F_L \rangle_{[1,6]}$. This analysis follows closely the study of ref. [6, 7] (see also refs. [9, 11, 12, 26, 54]). The aim is to reevaluate the constraints on the Wilson coefficients taking into account the following updates:

- Updated averages for $\langle A_{\text{FB}} \rangle_{[1,6]}^{\text{exp}}$ and $\langle F_L \rangle_{[1,6]}^{\text{exp}}$ including the new measurements in ref. [4] by the LHCb collaboration:

$$\langle A_{\text{FB}} \rangle_{[1,6]}^{\text{LHCb}} = -0.18^{+0.06+0.01}_{-0.06-0.02} \quad (4.1)$$

$$\langle F_L \rangle_{[1,6]}^{\text{LHCb}} = 0.66^{+0.06+0.04}_{-0.06-0.03} \quad (4.2)$$

The experimental averages for these two observables are collected in table 4.

- Updated theoretical predictions for $\langle A_{\text{FB}} \rangle_{[1,6]}$ and $\langle F_L \rangle_{[1,6]}$ including subleading corrections of order $|V_{ub}V_{us}|/|V_{tb}V_{ts}|$, as well as a recent update for f_B [51]:

$$f_B = 190 \pm 4 \text{ MeV} \tag{4.3}$$

- Analysis of the constraints using a consistent (frequentist) statistical approach detailed in appendix C.

These results will be used in the following sections to study the impact of the inclusion of binned observables, in view of (a) recent measurements at LHCb and (b) the impressive prospects for the near future measurements of q^2 -dependent $B \rightarrow K^*\ell^+\ell^-$ observables by the same collaboration.

We consider the following observables: $\text{BR}(B \rightarrow X_s\gamma)$, $S_{K^*\gamma}$, $A_I(B \rightarrow K^*\gamma)$, $\text{BR}(B \rightarrow X_s\mu^+\mu^-)$, $\langle A_{\text{FB}} \rangle_{[1,6]}$ and $\langle F_L \rangle_{[1,6]}$. The experimental situation is summarised in table 4, together with the SM predictions. As discussed in refs. [6, 7], these observables can be classified as class-I (dependence only on $\mathcal{C}_7^{(\prime)}$), class-II (dependence only on $\mathcal{C}_7^{(\prime)}$, $\mathcal{C}_9^{(\prime)}$, $\mathcal{C}_{10}^{(\prime)}$) and class-III (depending on all these plus other operators, e.g., scalar operators). The analysis is divided into four NP scenarios:

- *Scenario A*: New Physics in \mathcal{C}_7 and \mathcal{C}'_7 only, real values only.
- *Scenario B*: New Physics in \mathcal{C}_7 , \mathcal{C}'_7 , \mathcal{C}_9 , \mathcal{C}_{10} only, real values only.
- *Scenario B'*: New Physics in \mathcal{C}_7 , \mathcal{C}'_7 , \mathcal{C}'_9 , \mathcal{C}'_{10} only, real values only.
- *Scenario C*: New Physics in \mathcal{C}_7 , \mathcal{C}'_7 , $\mathcal{C}_9^{(\prime)}$, $\mathcal{C}_{10}^{(\prime)}$, real values only.

The constraints from $\text{BR}(B \rightarrow X_s\gamma)$, $S_{K^*\gamma}$, $A_I(B \rightarrow K^*\gamma)$ and $\text{BR}(B \rightarrow X_s\mu^+\mu^-)$ are implemented using the formulas presented in sections 2.3 and 2.5 of ref. [6]. The constraints are set on the shift of the Wilson coefficients with respect to their SM value at the hadronic scale μ_b . Concerning $\langle A_{\text{FB}} \rangle_{[1,6]}$ and $\langle F_L \rangle_{[1,6]}$, we provide the corresponding coefficients for the semi-numerical expressions of the integrated observables in appendix B, with an update of $\langle A_{\text{FB}} \rangle_{[1,6]}$ and $\langle F_L \rangle_{[1,6]}$ following the definition of the binned quantities described in section 2.

Inside the framework defined in ref. [6], the three observables $\text{BR}(B \rightarrow X_s\gamma)$, $S_{K^*\gamma}$ and $A_I(B \rightarrow K^*\gamma)$ (class-I observables) are insensitive to New Physics contributions to Wilson coefficients other than the electromagnetic operators $\mathcal{C}_7^{(\prime)}$. Therefore, the constraints from these (class-I) observables are common to all NP scenarios and affect only \mathcal{C}_7 and \mathcal{C}'_7 . The joint 68.3% and 95.5% C.L. constraints on \mathcal{C}_7 and \mathcal{C}'_7 are shown in the left panel of figure 4, together with the individual constraint from each observable. We find that the isospin asymmetry in $K^*\gamma$ disfavours the “flipped-sign” solution for \mathcal{C}_7 , and helps reducing the significance of the regions where $|\mathcal{C}'_7| \simeq \mathcal{C}_7^{\text{SM}}$, $\mathcal{C}_7 \simeq 0$. We recall that our scenarios assume the coefficients of the chromomagnetic operators $\mathcal{O}_8^{(\prime)}$ to be equal to their SM values.

We then consider the impact of adding the class-III observables ($\text{BR}(B \rightarrow X_s\mu^+\mu^-)$, $\langle A_{\text{FB}} \rangle_{[1,6]}$ and $\langle F_L \rangle_{[1,6]}$) in the different NP scenarios in turn. Within Scenario A, we obtain the joint 68.3% and 95.5% C.L. constraints in the \mathcal{C}_7 - \mathcal{C}'_7 plane shown in the right panel of

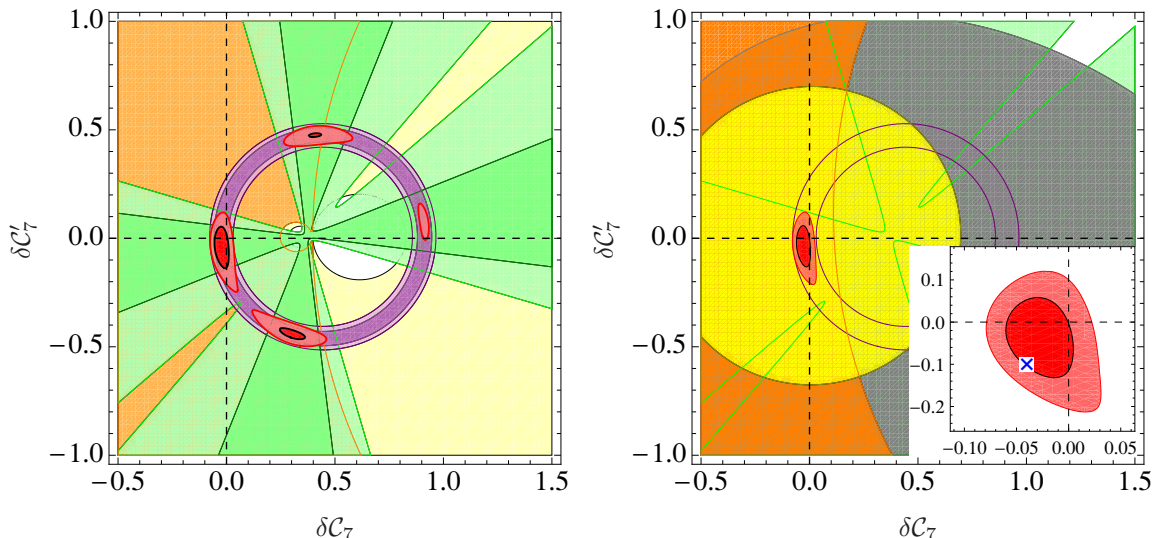


Figure 4. 68.3% (dark) and 95.5% (light) C.L. constraints on $\delta\mathcal{C}_7(\mu_b)$, $\delta\mathcal{C}'_7(\mu_b)$. Left: Class-I observables — $\text{BR}(B \rightarrow X_s \gamma)$ (purple), $S_{K^* \gamma}$ (green) and $A_I(B \rightarrow K^* \gamma)$ (yellow/orange). Right: Scenario A (class-I and class-III) — $\text{BR}(B \rightarrow X_s \mu^+ \mu^-)$ (yellow), $\langle A_{\text{FB}} \rangle_{[1,6]}$ (orange) and $\langle F_L \rangle_{[1,6]}$ (gray). The combined constraints are shown in red. The cross indicates the position of the benchmark point a . The origin $(0,0)$ corresponds to the SM point.

Point	$\delta\mathcal{C}_7(\mu_b)$	$\delta\mathcal{C}'_7(\mu_b)$	$\delta\mathcal{C}_9(\mu_b)$	$\delta\mathcal{C}_{10}(\mu_b)$	$\delta\mathcal{C}'_9(\mu_b)$	$\delta\mathcal{C}'_{10}(\mu_b)$
a	-0.04	-0.1	0	0	0	0
$b1$	-0.03	-0.09	-1.5	-1	0	0
$b2$	0.3	-0.4	1	6	0	0
$b3$	0.45	0.45	-9	2	0	0
$b4$	0.9	0.1	-9	8	0	0
b'	-0.05	-0.15	0	0	3	1.5
c	0.4	-0.45	-7	5	3	5

Table 5. NP benchmark points used in the analysis of section 5.

figure 4. In this case, the $B \rightarrow K^* \mu^+ \mu^-$ forward-backward asymmetry disfavors strongly the two regions with large $\delta\mathcal{C}'_7$ allowed at the 95.5% C.L. by class-I observables. This plot features also the benchmark point a , defined in table 5 and used in the next section to study the power of q^2 -dependent $B \rightarrow K^* \ell^+ \ell^-$ observables to discriminate among NP scenarios.

The joint constraints within Scenario B are shown in figure 5. Also shown is the constraint from $B_s \rightarrow \mu^+ \mu^-$, which is a direct constraint on \mathcal{C}_{10} . We see that at the 95.5% C.L., there are four allowed regions in space of Wilson coefficients. The four benchmark points $b1$ - $b4$ are also indicated, one in each of the four allowed regions. Finally, in figures 6 and 7 we show the constraints within Scenarios B' and C, together with the corresponding benchmark points b' and c' . The $B_s \rightarrow \mu^+ \mu^-$ branching ratio constrains \mathcal{C}'_{10} directly in Scenario B', but only $|\mathcal{C}_{10} - \mathcal{C}'_{10}|$ in Scenario C. We plot the constraint based on ref. [12],

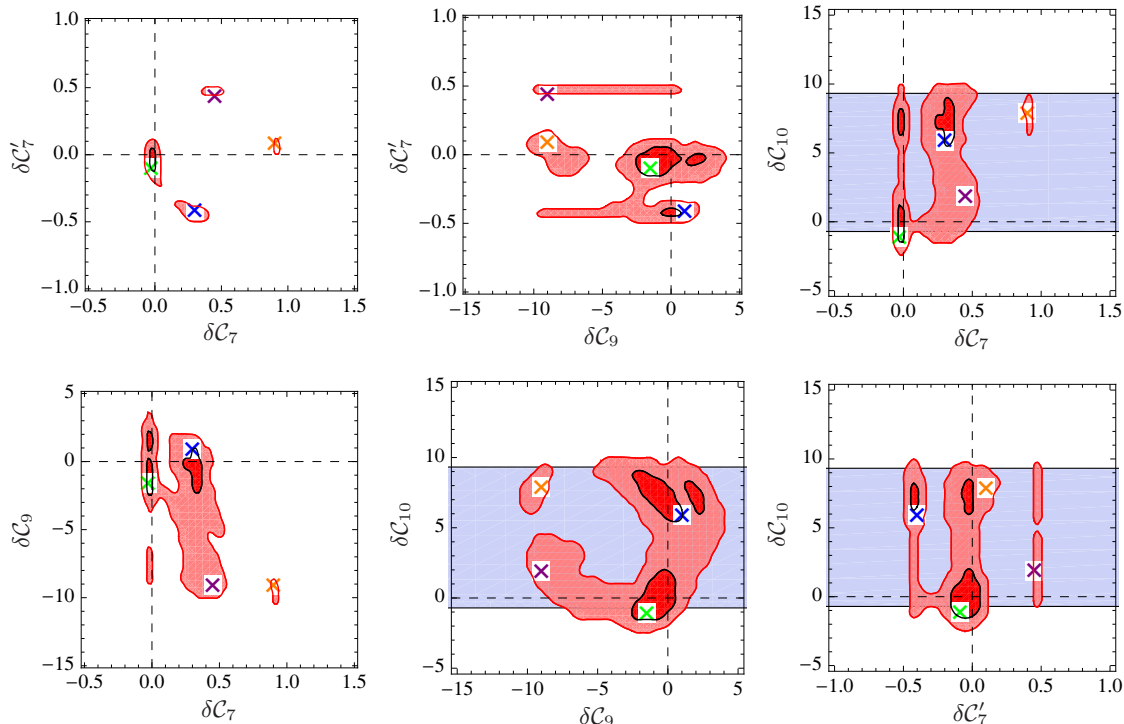


Figure 5. 68.3% (light red) and 95.5% (dark red) C.L. contours for $\delta\mathcal{C}_7^{(\prime)}(\mu_b)$, $\delta\mathcal{C}_9(\mu_b)$, $\delta\mathcal{C}_{10}(\mu_b)$ in Scenario B. The crosses indicate benchmark points b_1 (green), b_2 (blue), b_3 (purple) and b_4 (orange). The blue band corresponds to the $B_s \rightarrow \mu^+\mu^-$ constraint.

including the $O(\Delta\Gamma_s)$ correction needed to connect theory and experiment [21–23]. (Using the value from ref. [19] would result in a slightly tighter constraint. Since there is only an experimental bound on this branching ratio, and since the status of the theoretical prediction is unclear, we refrain from including this piece of information in our combined constraints on the Wilson coefficients.)

As a final comment, we note that the “flipped-sign” solution for \mathcal{C}_7 is in general disfavoured, but not very significantly. In fact, in Scenario B, due to the loosening of the constraint from $\langle A_{\text{FB}} \rangle_{[1,6]}$ when \mathcal{C}_9 , \mathcal{C}_{10} are allowed not to vanish, this flipped-sign solution reappears inside the 95.5% C.L. region. More precise constraints from $A_I(B \rightarrow K^*\gamma)$ and $\langle A_{\text{FB}} \rangle_{[1,6]}$ should help to settle this question.

5 Complementarity of observables for NP studies

The complementarity of the different angular observables in the identification of possible NP effects is manifest even if we deal exclusively with observables integrated over the whole $[1, 6] \text{ GeV}^2$ ranges. For example, New Physics contributing predominantly to \mathcal{C}_{10} will most likely push substantially $\langle P_4 \rangle_{[1,6]}$ below its SM value⁵ (see ref. [32]). However, this will be

⁵To be specific, since this discussion is for illustrative purposes only, we focus on the New Physics scenarios and the observables $P_{4,5,6}$ of ref. [32]. Below we will study specific benchmark points and the primed observables $P'_{4,5,6}$.

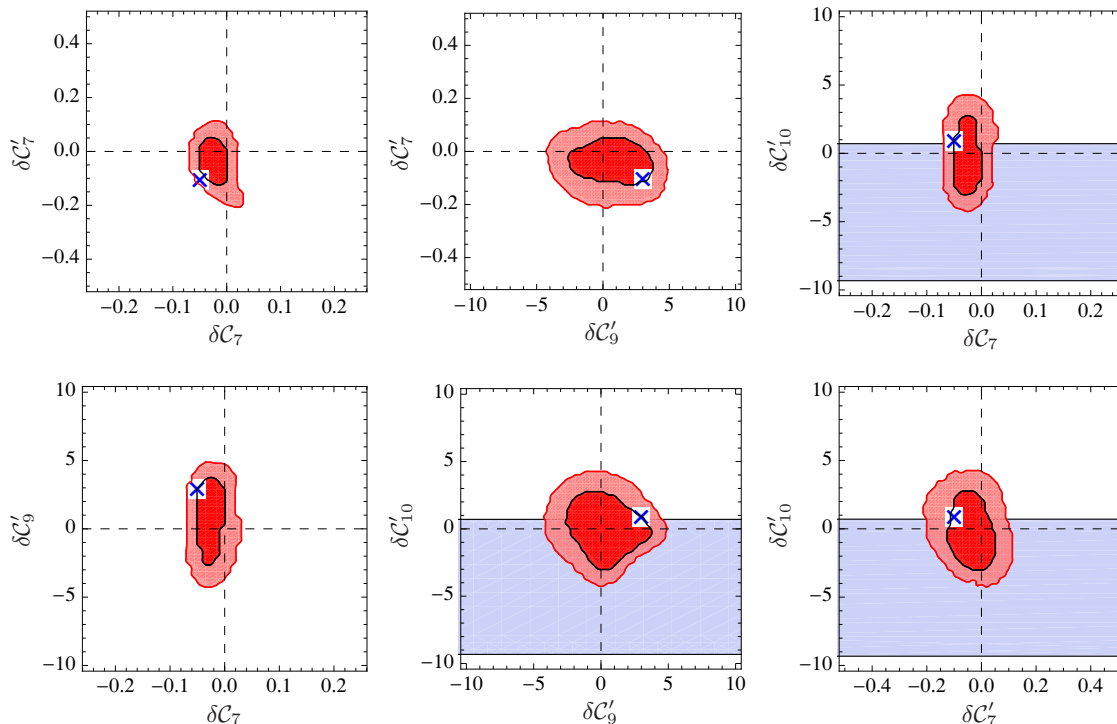


Figure 6. 68.3% (light red) and 95.5% (dark red) C.L. contours for $\delta\mathcal{C}_7^{(\prime)}(\mu_b)$, $\delta\mathcal{C}_9^{(\prime)}(\mu_b)$, $\delta\mathcal{C}_{10}^{(\prime)}(\mu_b)$ in Scenario B'. The cross is benchmark point b' . The $B_s \rightarrow \mu^+\mu^-$ constraint is indicated as a blue band.

essentially indistinguishable from a NP contribution to \mathcal{C}'_{10} , which has a very similar effect on P_4 . This ambiguity can be resolved looking at the measured value of $\langle P_1 \rangle_{[1,6]}$, because a New Physics contribution to \mathcal{C}'_{10} can enhance this observable considerably above its SM prediction, while \mathcal{C}_{10} has a negligible effect.

Another example would correspond to a moderate enhancement of $\langle P_1 \rangle_{[1,6]}$. Assuming no significant deviation is seen in $\langle P_4 \rangle_{[1,6]}$, this could signal a non-SM value of \mathcal{C}'_9 or $\mathcal{C}_7, \mathcal{C}'_7$ (in the island around the SM point in the \mathcal{C}_7 - \mathcal{C}'_7 plane). The former case will not give any substantial deviation in $\langle P_5 \rangle_{[1,6]}$, while the latter tends to increase $\langle P_5 \rangle_{[1,6]}$ above its SM value. Comparing $\langle P_1 \rangle_{[1,6]}$, $\langle P_4 \rangle_{[1,6]}$ and $\langle P_5 \rangle_{[1,6]}$ would thus help to distinguish among these scenarios.

The discriminating power of the observables is substantially increased when we consider, not only different angular observables, but also their q^2 -dependence. A New Physics contribution to \mathcal{C}'_{10} (in the scenario considered above) will generally increase $\langle P_1 \rangle_{[1,6]}$, but its effect on $\langle P_4 \rangle_{[1,6]}$ and $\langle P_5 \rangle_{[1,6]}$ could be washed out by simultaneous contributions to \mathcal{C}_9 and \mathcal{C}_{10} . However, if this enhancement of $\langle P_1 \rangle_{[1,6]}$ is also present in the region $\sim 4-6 \text{ GeV}^2$ (for instance with $\langle P_1 \rangle_{[4.3,6]}$ shifted upwards with respect to the SM), the NP effect cannot be misinterpreted as a non-SM value of $\mathcal{C}_7^{(\prime)}$ or \mathcal{C}'_9 , because these would only enhance P_1 in the region $\sim 1-3 \text{ GeV}^2$.

These considerations apply equally well in terms of *constraints* when no deviation from the SM is observed. In section 8 we will see explicitly how different observables

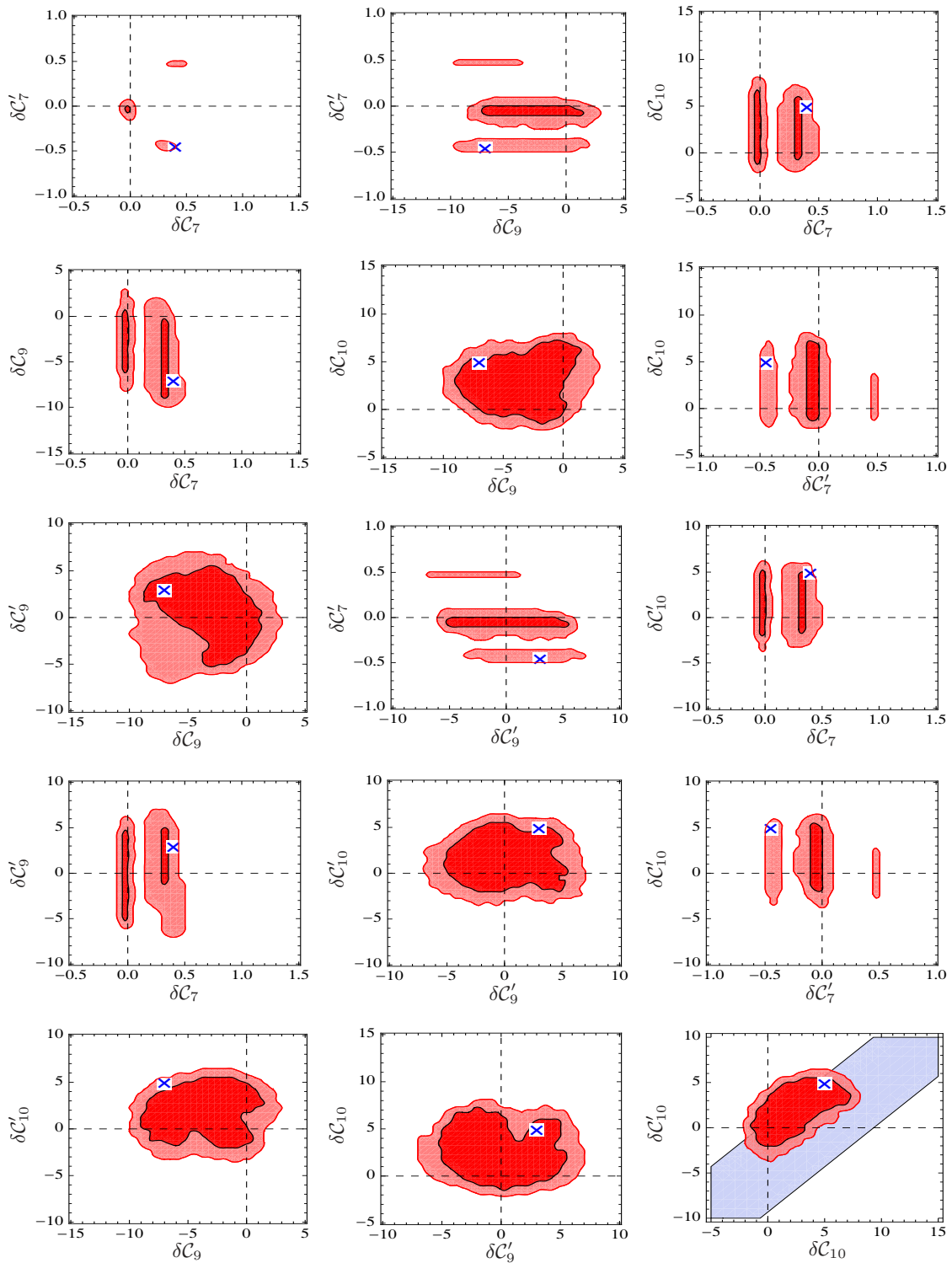


Figure 7. 68.3% (light red) and 95.5% (dark red) C.L. contours for $\delta\mathcal{C}_7(\mu_b)$, $\delta\mathcal{C}'_7(\mu_b)$, $\delta\mathcal{C}'_9(\mu_b)$, $\delta\mathcal{C}'_{10}(\mu_b)$ in Scenario C. The cross is benchmark point c . The $B_s \rightarrow \mu^+\mu^-$ constraint is indicated as a blue band.

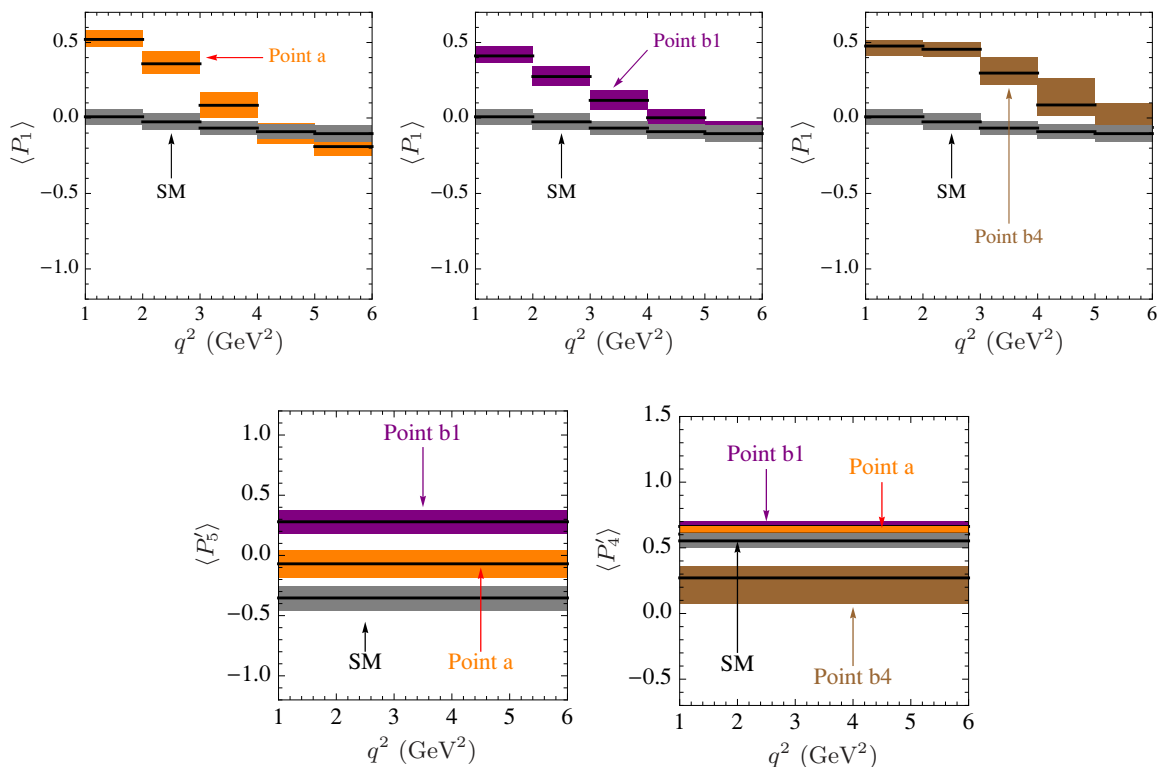


Figure 8. Comparison of SM predictions for $\langle P_1 \rangle$, $\langle P'_4 \rangle_{[1,6]}$, $\langle P'_5 \rangle_{[1,6]}$ and the predictions within benchmark points a , $b1$, $b4$ given in table 5.

constrain the NP contributions to the Wilson coefficients. It is already useful to build an intuition on how different observables in different q^2 -bins are affected by shifts in the different Wilson coefficients, in order to have a better idea of the most promising observables in each case. Here we study briefly the effect of different New Physics scenarios on the integrated observables studied in section 3. We focus on a set of “benchmark points” consistent with $\text{BR}(B \rightarrow X_s \gamma)$, $S_{K^* \gamma}$, $A_I(B \rightarrow K^* \gamma)$, $\text{BR}(B \rightarrow X_s \mu^+ \mu^-)$, $\langle A_{\text{FB}} \rangle_{[1,6]}$ and $\langle F_L \rangle_{[1,6]}$, according to the analysis of section 4. These points are specified in table 5, and also indicated in figures 4-7.

Each NP benchmark point gives a prediction for each observable in each q^2 -bin. In figures 8, 9 and 10 we show the most significant cases, exemplifying the previous discussion. In these figures, the gray rectangles correspond to the SM binned predictions, corresponding to figures 1 and 2. The colored rectangles correspond to the NP scenarios represented by each benchmark point collected in table 5. We can draw a few conclusions:

- $\langle P_1 \rangle$ in the region $q^2 \lesssim 3 \text{ GeV}^2$ is able to discriminate the points a , $b1$ and $b4$ from the SM, but gives similar predictions for these scenarios. These scenarios could be resolved by $\langle P'_4 \rangle_{[1,6]}$ and $\langle P'_5 \rangle_{[1,6]}$ (see figure 8).
- The points $b2$, $b3$ and b' enhance the observables $\langle P_1 \rangle$ and $\langle P'_4 \rangle$ substantially, and the impact on $\langle P_1 \rangle$ ($\langle P'_4 \rangle$) is more important for $q^2 > 3 \text{ GeV}^2$ ($q^2 < 3 \text{ GeV}^2$) (see figure 9).

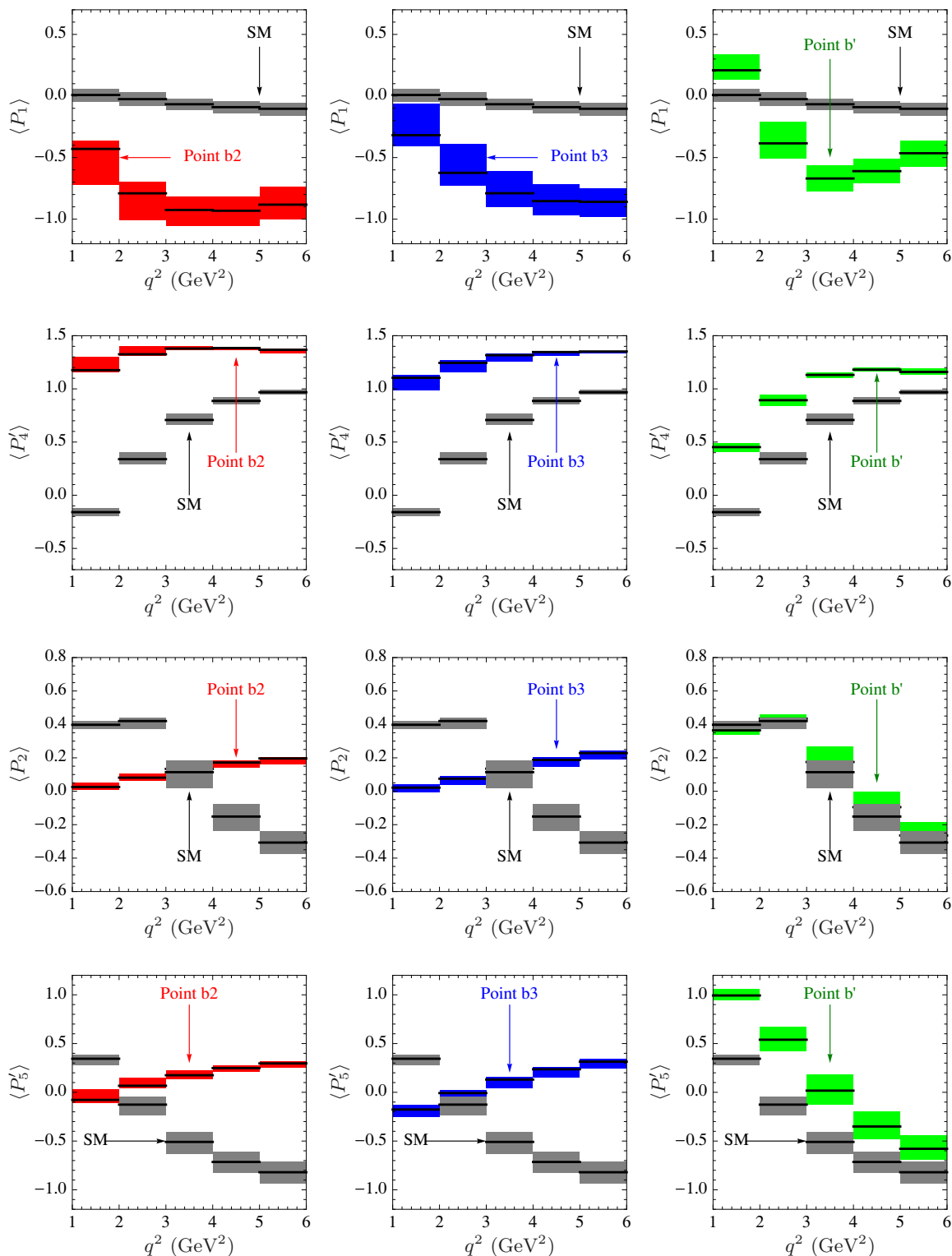


Figure 9. Comparison of SM predictions for $\langle P_1 \rangle$, $\langle P_2 \rangle$, $\langle P'_4 \rangle$, $\langle P'_5 \rangle$ and the predictions within benchmark points b_2 , b_3 , b' given in table 5.

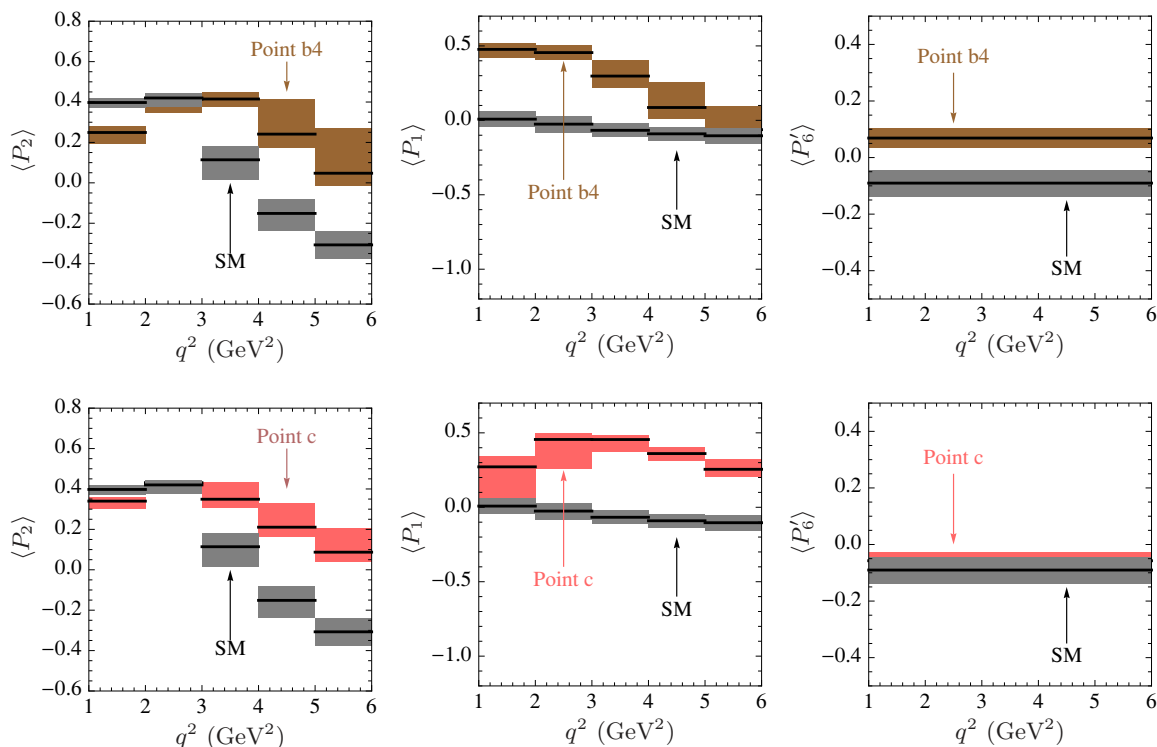


Figure 10. Comparison of SM predictions for $\langle P_1 \rangle$, $\langle P_2 \rangle$, $\langle P'_6 \rangle$ and the predictions within benchmark points b_4 , c given in table 5.

- $\langle P_1 \rangle$ and $\langle P'_4 \rangle$ do not allow one to discriminate among b_2 , b_3 and b' , but $\langle P_2 \rangle$ and $\langle P'_5 \rangle$ exhibit distinctive behaviours for b_2 and b_3 (figure 9). For instance, a suppression of $\langle P_1 \rangle$ below ~ -0.5 together with an enhancement of $\langle P_2 \rangle_{[4,6]}$ above ~ 0 would favour b_2 and b_3 with respect to b' . This conclusion could be verified by a suppression of $\langle P_2 \rangle_{[1,3]}$ below its SM value and an enhancement of $\langle P'_5 \rangle_{[4,6]}$.
- A similar situation occurs with the points b_4 and c . The observable $\langle P_2 \rangle_{[5,6]}$ could favour these scenarios, but cannot distinguish among them. However, $\langle P_1 \rangle_{[1,2]}$ and $\langle P_1 \rangle_{[5,6]}$ can discriminate these scenarios, as well as $\langle P'_6 \rangle_{[1,6]}$ if the experimental values are accurate enough (see figure 10).

A full set of predictions for all benchmark points in comparison with the SM predictions can be found in figures 8 and 9 of ref. [57].

6 The benefit of using clean observables

In this section we discuss the advantages of using clean observables in analyses of $B \rightarrow K^* \ell^+ \ell^-$ as opposed to other observables such as S_3 or A_{im} . For definiteness we focus on P_1 and S_3 , but it should be kept in mind that the conclusions are more general.

As discussed extensively in refs. [30–33], clean observables are constructed in such a way that at LO and at large recoil, an exact cancellation of the form-factor dependence

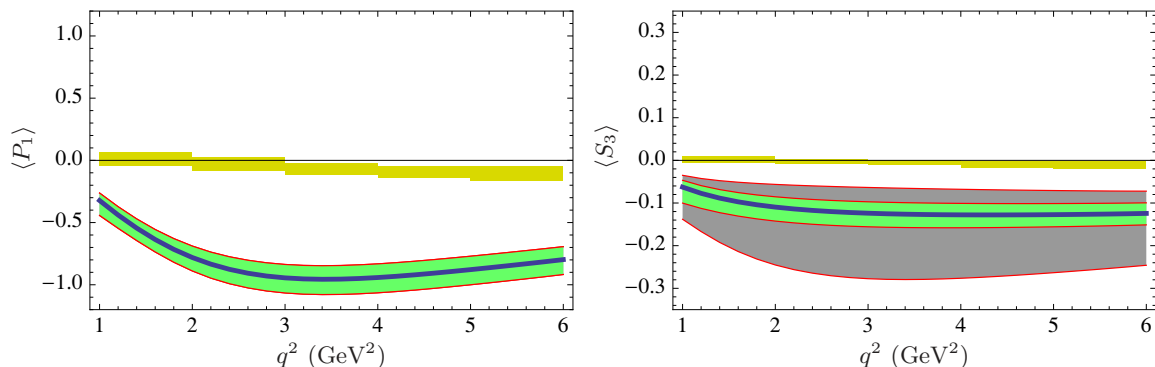


Figure 11. Predictions in the SM and in the case of NP at the benchmark point $b2$ for P_1 (left) and S_3 (right). The yellow boxes are the SM predictions integrated in five 1 GeV^2 bins. The blue curve corresponds to the central values for the NP scenario. The green band is the total uncertainty considering the form factors of refs. [26, 28], while the gray band is the total uncertainty obtained using the form factors of ref. [27]. In the case of P_1 the gray band is barely visible.

occurs. This indicates that clean observables should be stable under variation of hadronic uncertainties, as opposed to other observables, such as F_L , A_{FB} , S_3 , etc. This is relevant because of the spread of published errors in the determination of form factors from light-cone sum rules (see refs. [27, 28] and the introduction).

If the form factors of ref. [27] are used in the evaluation of F_L , for example, the error bars get enlarged by a factor of three. On the contrary, this enlargement does not happen in the case of P_1 , which is practically insensitive to these uncertainties. In the case of S_3 , an accidental circumstance makes its SM uncertainty smaller than what one would infer from the fact that $S_3 \sim P_1 F_T$ (that is, a similar percentual enhancement of the errors as F_L). The fact that $P_1 \sim 0$ in the SM, makes S_3 almost insensitive to the error in F_L only near the SM point. This makes the situation with S_3 a bit more subtle. The important point here is that in the presence of New Physics, an enhancement of P_1 produces an enlargement of the error bars in the theoretical prediction for S_3 automatically, which makes S_3 almost unable of discriminating between NP models where P_1 does not vanish.

In figure 11 we show the SM predictions and the predictions for benchmark point $b2$ (see table 5) for P_1 and S_3 calculated with both choices of form factors (refs. [27] and [28]). We find that:

- The SM prediction for P_1 is insensitive to the choice of form factors.
- The SM prediction for S_3 shows a moderate dependence on the choice of form factors, and hadronic uncertainties are enlarged up to a 50% when using the form factors of ref. [27] compared to those in ref. [28].
- The NP prediction for P_1 is insensitive to the choice of form factors.
- The NP prediction for S_3 is *very* sensitive to the choice of form factors. Indeed, the hadronic uncertainties increase from a factor 2 to a factor 3 when using the form factors of ref. [27].

The conclusion is that, in New Physics analyses of $B \rightarrow K^* \ell^+ \ell^-$, when using the observable S_3 , one is obliged to take into account hadronic uncertainties at each point in the NP parameter space, and these uncertainties can be substantial. On the other hand, the observable P_1 is not affected by this disease, and it is therefore a much more robust observable. Moreover, from figure 11 we can see that P_1 has the potentiality to exclude completely a NP scenario given by benchmark point $b2$, while S_3 can barely distinguish this point from the SM case at more than 2σ . We expect the same results to hold in other regions of the NP parameter space. The conclusion does not change when we consider the binned observables. For example, the corresponding predictions at benchmark point $b2$ for $\langle P_1 \rangle$ and $\langle S_3 \rangle$ in the bin $[1, 6] \text{ GeV}^2$ are

$$\langle P_1 \rangle_{[1,6]}^{b2} = -0.82 \pm 0.10, \quad \langle S_3 \rangle_{[1,6]}^{b2} = -0.16 \pm 0.08. \quad (6.1)$$

We would like to emphasise that similar conclusions are expected for other form factor dependent observables. For example, the observable $A_{\text{im}} \sim S_9 \sim P_3 F_T$ will also be protected from large hadronic uncertainties near the SM point, since $P_3 \sim 0$ in the SM. However, complex NP can enhance considerably P_3 while being consistent with other data (see ref. [32]). In these NP scenarios, A_{im} is expected to suffer from a similar problem as the one described for S_3 (a problem that does not affect P_3). Furthermore, such problems may also happen also in the case of CP conjugated observables such as A_9 and A_3 .

In view of this situation, one should be particularly careful with a rather usual hypothesis consisting in assigning the same hadronic uncertainty on observables in any NP scenario as in the SM case. This might be a good approximation in the case of clean observables (see for instance the left-hand side of figure 11, where the uncertainty on P_1 is indeed similar from the SM case to the NP scenario considered), but it can be very misleading for other quantities, sensitive to form factors (as illustrated on the right-hand side of figure 11, where the uncertainty on S_3 is significantly enlarged from the SM value to the NP scenario considered).

7 Extracting clean observables from existing experimental measurements

As pointed out in section 6, there is a clear advantage in using the clean observables $P_{1,2,3}$ rather than S_3 , A_{im} and A_{FB} (or F_L) to put constraints on the Wilson coefficients. However, this can be achieved if the experimental fits are performed consistently considering these observables (see ref. [32]). In its latest experimental analysis of q^2 -dependent observables in $B \rightarrow K^* \ell^+ \ell^-$, the LHCb collaboration [4] has preferred to fit directly the coefficients of the angular distributions, providing only observables proportional to the J_i coefficients (in particular $\langle S_3 \rangle$ and $\langle A_{\text{im}} \rangle$). We have collected for reference the relevant experimental results for these observables in table 6.

Of course, we can compute the clean observables from the measurements provided using the formulas of sections 3 and 4 of ref. [32], or equivalently from eqs. (2.4). If this is done without knowing the correlation matrix (which is not provided by the experimental collaborations yet), one obtains errors much larger than the real uncertainties. It is still

	[2, 4.3]	[4.3, 8.68]	[1, 6]
$\langle A_{\text{FB}} \rangle$	$-0.20^{+0.08+0.01}_{-0.07-0.03}$	$0.16^{+0.05+0.01}_{-0.05-0.01}$	$-0.18^{+0.06+0.01}_{-0.06-0.02}$
$\langle F_L \rangle$	$0.74^{+0.09+0.02}_{-0.08-0.04}$	$0.57^{+0.05+0.04}_{-0.05-0.03}$	$0.66^{+0.06+0.04}_{-0.06-0.03}$
$\langle A_{\text{im}} \rangle$	$-0.02^{+0.10+0.05}_{-0.06-0.01}$	$0.02^{+0.07+0.01}_{-0.07-0.01}$	$0.07^{+0.07+0.02}_{-0.07-0.01}$
$\langle S_3 \rangle$	$-0.05^{+0.18+0.05}_{-0.12-0.01}$	$0.18^{+0.13+0.01}_{-0.13-0.01}$	$0.10^{+0.15+0.02}_{-0.16-0.01}$

Table 6. LHCb experimental results for binned observables (from ref. [4]).

Observable	Experiment	SM prediction
$\langle P_1 \rangle_{[2,4.3]}$	-0.19 ± 0.58	-0.051 ± 0.050
$\langle P_1 \rangle_{[4.3,8.68]}$	0.42 ± 0.31	-0.115 ± 0.060
$\langle P_1 \rangle_{[1,6]}$	0.29 ± 0.47	-0.055 ± 0.051
$\langle P_2 \rangle_{[2,4.3]}$	0.51 ± 0.27	0.227 ± 0.070
$\langle P_2 \rangle_{[4.3,8.68]}$	-0.25 ± 0.08	-0.422 ± 0.074
$\langle P_2 \rangle_{[1,6]}$	0.35 ± 0.14	0.080 ± 0.067
$\langle P_3 \rangle_{[2,4.3]}$	0.08 ± 0.35	-0.004 ± 0.024
$\langle P_3 \rangle_{[4.3,8.68]}$	-0.05 ± 0.16	-0.005 ± 0.027
$\langle P_3 \rangle_{[1,6]}$	-0.21 ± 0.21	-0.003 ± 0.024

Table 7. Experimental values for the clean observables P_1 , P_2 and P_3 within different q^2 -bins, extracted from the measurements of S_3 , A_{im} , A_{FB} and F_L , and their SM predictions.

worth exploring the current situation based on these observables while waiting for correlated values. In addition, higher-statistics analyses from LHCb are expected to reduce the experimental errors on these observables considerably very soon.

Attending to these considerations, experimental values for $P_{1,2,3}$ can be derived from the measurements of S_3 , A_{im} , A_{FB} and F_L in table 6 by means of eq. (2.4):

$$\langle P_1 \rangle_{\text{bin}} = \frac{2\langle S_3 \rangle_{\text{bin}}}{1 - \langle F_L \rangle_{\text{bin}}}, \quad \langle P_2 \rangle_{\text{bin}} = -\frac{2}{3} \frac{\langle A_{\text{FB}} \rangle_{\text{bin}}}{(1 - \langle F_L \rangle_{\text{bin}})}, \quad \langle P_3 \rangle_{\text{bin}} = -\frac{\langle A_{\text{im}} \rangle_{\text{bin}}}{(1 - \langle F_L \rangle_{\text{bin}})}. \quad (7.1)$$

In table 7 we present the resulting experimental values for $\langle P_1 \rangle$, $\langle P_2 \rangle$, $\langle P_3 \rangle$ in the different bins, together with their SM predictions. The experimental errors are calculated in the following way. We first add in quadrature both errors in table 6, and symmetrise upper and lower uncertainties. Assuming these errors are Gaussian, the errors for $P_{1,2,3}$ are obtained by the usual error propagation formula from eqs. (7.1).

8 Present and future constraints from q^2 -dependent $B \rightarrow K^* \ell^+ \ell^-$ observables

The recent LHCb measurements for q^2 -dependent $B \rightarrow K^* \ell^+ \ell^-$ observables [4] are divided into four bins — if we restrict ourselves to the low- q^2 region. These bins are $[0.05, 2]$,

[2, 4.3], [4.3, 8.68] and [1, 6] GeV². These results yield a first glimpse of the future, where precise measurements of the full angular distribution within fine q^2 -bins will be available. The purpose of this section is to study the impact of the q^2 -dependent observables provided in ref. [4] on the constraints on the Wilson coefficients, and to analyze what are the future expectations concerning the constraints from these observables.

A brief discussion is in order concerning the results in ref. [4]. From the theory point of view the first bin [0.05, 2] is very difficult to control, since the decay rate contains contributions from light resonances below $q^2 \sim 1$ GeV². The third bin [4.3, 8.68] is also more difficult to handle theoretically, as it gets near to the $c\bar{c}$ resonance region. For this first attempt, we choose to drop the first bin and to consider constraints from the two others [2, 4.3], [4.3, 8.68]. We do not include [1, 6] as the results are likely to be strongly correlated with the two smaller bins considered for this study (we remind that the averaged experimental results in the bin [1, 6] have already been considered in section 4 in the case of $\langle A_{\text{FB}} \rangle$ and $\langle F_L \rangle$).

8.1 Constraints from binned $\langle A_{\text{FB}} \rangle$ and $\langle F_L \rangle$

We first consider the observables $\langle A_{\text{FB}} \rangle_{[2,4.30]}$, $\langle A_{\text{FB}} \rangle_{[4.30,8.68]}$, $\langle F_L \rangle_{[2,4.30]}$ and $\langle F_L \rangle_{[4.30,8.68]}$. The experimental numbers for these observables are given in table 6. We study separately the constraints derived using the set of form factors of ref. [26] or the ones from ref. [27], the later being the choice throughout this article. In this case the constraints are implemented using eq. (B.1) in appendix B together with the coefficients in tables 8 and 9. The constraints in the first case (form factors from ref. [26]) can be implemented from appendix B in ref. [57].

The individual 68.3% C.L. and combined 68.3% and 95.5% C.L. contours for these observables in the \mathcal{C}_7 - \mathcal{C}'_7 plane (Scenario A) are shown in figure 12. We see that:

- The constraints from $\langle A_{\text{FB}} \rangle$ are consistent with the SM at 95.5% C.L. Using the form factors in ref. [26], some tension is caused by $\langle A_{\text{FB}} \rangle_{[2,4.30]}$, while $\langle A_{\text{FB}} \rangle_{[4.30,8.68]}$ is compatible with the SM at 68.3% C.L. This tension disappears if the form factors of ref. [27] are used in the SM predictions. In this case the 95.5% C.L. region widens considerably.
- With form factors of ref. [26], the constraints from $\langle F_L \rangle$ show a discrepancy with the SM, which is just outside the 95.5% C.L. region. Indeed $\langle F_L \rangle_{[4.30,8.68]}$ has a clear tendency to avoid the SM point. This tension disappears completely if the form factors of ref. [27] are used for the SM predictions.
- Committing oneself to form factors in ref. [26] and taking seriously these tensions would require a NP that affects simultaneously A_{FB} around $q^2 \sim 3$ GeV² and F_L around $q^2 \sim 6$ GeV². Predictions derived from form factors in ref. [27] are perfectly consistent with the SM; this is the conservative conclusion.

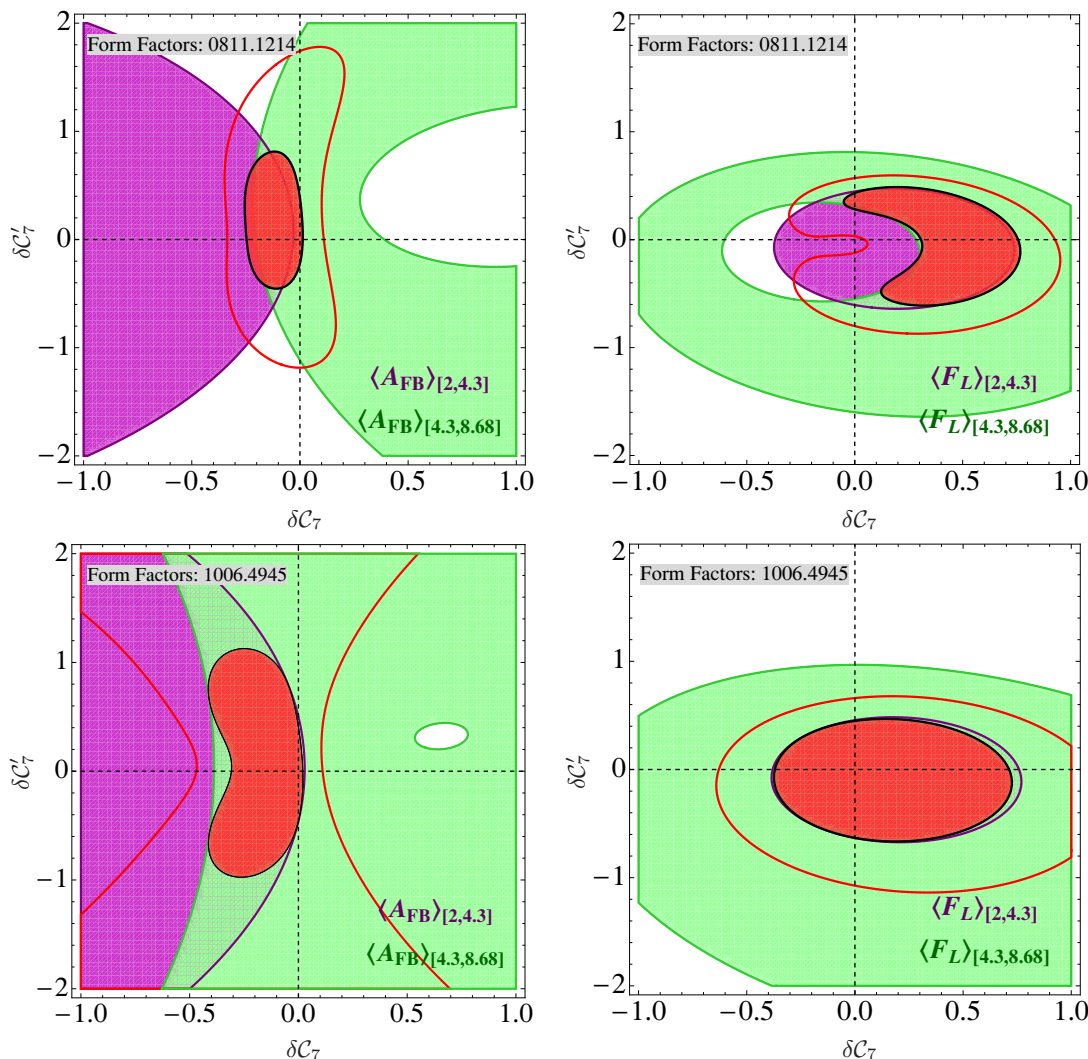


Figure 12. Individual 68.3% C.L. constraints in the $\delta\mathcal{C}_7(\mu_b)$ - $\delta\mathcal{C}'_7(\mu_b)$ plane from $\langle A_{\text{FB}} \rangle_{[2,4,3]}$ and $\langle A_{\text{FB}} \rangle_{[4,30,8,68]}$ (left), and from $\langle F_L \rangle_{[2,4,3]}$ and $\langle F_L \rangle_{[4,30,8,68]}$ (right), taking form factors of ref. [26] (up) or ref. [27] (down). The combined 68.3% C.L. (red filled) and 95.5% C.L. (red contour) regions are also shown. The origin corresponds to the SM value.

8.2 Constraints from binned $\langle P_1 \rangle$, $\langle P_2 \rangle$ and $\langle P_3 \rangle$

We now consider the constraints from the observables $\langle P_i \rangle_{[2,4,30]}$, $\langle P_i \rangle_{[4,30,8,68]}$, $\langle P_i \rangle_{[1,6]}$ with $i = 1, 2, 3$. The experimental values and SM predictions for these observables are collected in table 7. The constraints are implemented using eq. (B.1) in appendix B together with the coefficients in tables 10, 11 and 12.

The individual constraints from these observables in the \mathcal{C}_7 - \mathcal{C}'_7 plane (Scenario A) are shown in figure 13. We see that:

- The constraints from $\langle P_1 \rangle$ are not very stringent yet. However there is a very mild discrepancy of $\langle P_1 \rangle_{[4,30,8,68]}$ with respect to the SM, as well as the combined constraint from the two bins. This result is not affected by form factor uncertainties.

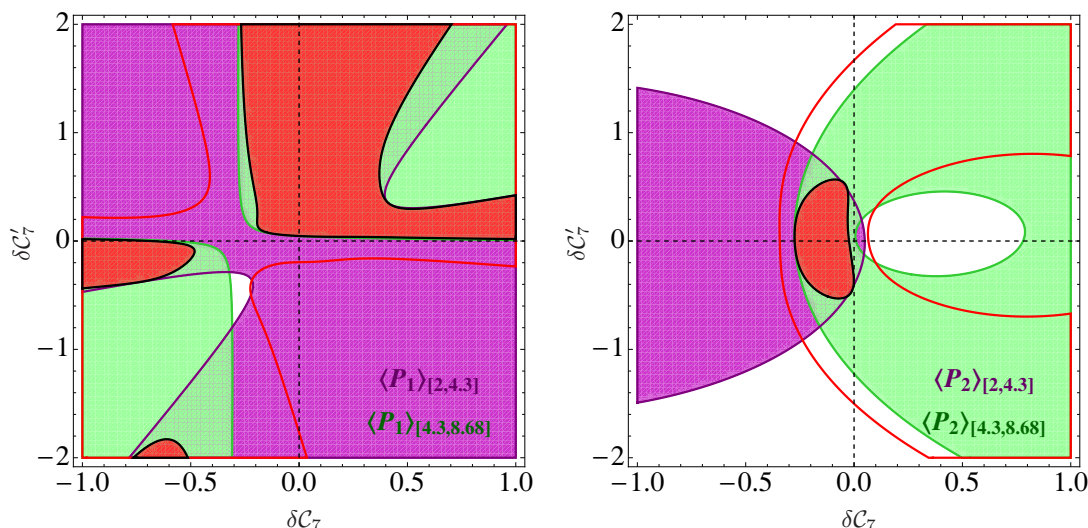


Figure 13. Left: individual 68.3% C.L. constraints in the $\delta\mathcal{C}_7(\mu_b)$ - $\delta\mathcal{C}'_7(\mu_b)$ plane from the integrated clean observables $\langle P_1 \rangle_{[2,4,30]}$ and $\langle P_1 \rangle_{[4.30,8.68]}$, together with the combined result. The red region and contour correspond to the combined 68.3% and 95.5% C.L. regions. Right: same analysis for $\langle P_2 \rangle_{\text{bin}}$. The origin corresponds to the SM point. Currently, $\langle P_3 \rangle_{\text{bin}}$ does not provide any constraint on $\mathcal{C}_7(\mu_b)$ and $\mathcal{C}'_7(\mu_b)$.

- The constraints from $\langle P_2 \rangle$ are already quite interesting. The two bins point towards a negative value of $\delta\mathcal{C}_7$, and the SM point is just outside the 68.3% C.L. region. Again, this result is not affected by form factor uncertainties. While the theoretical prediction for $\langle P_2 \rangle_{[4.30,8.68]}$ can suffer from the proximity of the bin to the $c\bar{c}$ resonance, we point out that the same tendency to negative $\delta\mathcal{C}_7$ is hinted at by the observable $\langle P_2 \rangle_{[1,6]}$, indicating that this is not a feature introduced by the data above 6 GeV². An enhancement of $\langle P_2 \rangle$ in the full low- q^2 region would be consistent with NP scenarios b_4 and c (see figure 10).
- The constraints from $\langle P_3 \rangle$ are inconclusive for the time being. This could be guessed already from the NP analysis of section 5 (see figure 8 in ref. [57]). It is well known that the CP-averaged version of $\langle P_3 \rangle_{\text{bin}}$ (the one we are considering here) is not very sensitive to NP, and that the corresponding CP-asymmetry might be more interesting when constraining NP (see for example ref. [12]). In this case we would suggest to focus, instead of A_9 , on the corresponding clean CP-asymmetry, since A_9 can be affected by the problems discussed in section 6 concerning its sensitivity to form factors.

We stress again that these constraints should be considered as conservative, since they are based on the experimental numbers extracted in section 7 in absence of experimental correlations. Therefore, the uncertainties of $P_{1,2,3}$ that we quote in table 7 are most probably overestimated, and could be reduced significantly once experimental correlations are available.

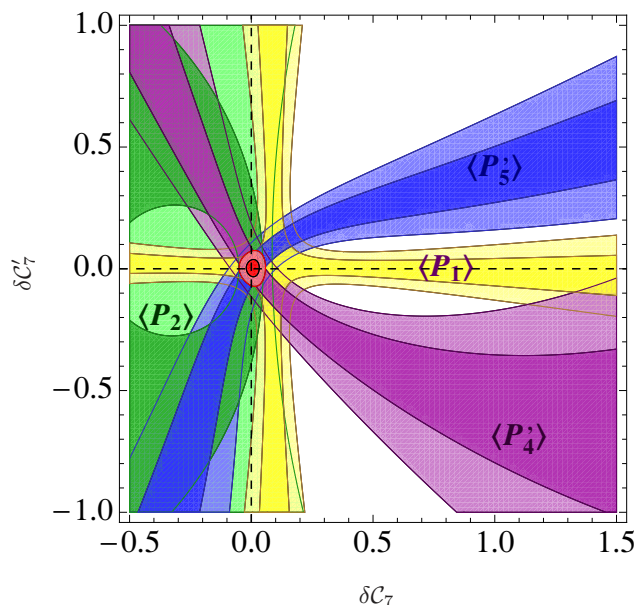


Figure 14. Individual constraints in the $\delta\mathcal{C}_7$ - $\delta\mathcal{C}'_7$ plane from hypothetical measurements of the observables $\langle P_1 \rangle_{[2,4,3]}$, $\langle P_2 \rangle_{[2,4,3]}$, $\langle P'_4 \rangle_{[2,4,3]}$ and $\langle P'_5 \rangle_{[2,4,3]}$, corresponding to central values equal to the SM predictions and an experimental uncertainty $\sigma_{\text{exp}} = 0.10$. The combined 68.3% (dark red) and 95.5% (light red) C.L. regions are also shown.

8.3 Future prospects

The experimental numbers in table 6 for the various observables, as measured by the LHCb collaboration, contain uncertainties at the level of ~ 0.10 , and up to ~ 0.20 for the observable $\langle S_3 \rangle$. The numbers for $\langle P_{1,2,3} \rangle$ extracted in the previous section from the measurements in table 6 contain larger uncertainties, up to ~ 0.5 (see table 7). As discussed above, these errors are probably overestimated since they do not take into account the relevant correlations among the observables. It is reasonable to expect that a direct extraction of $\langle P_{1,2,3} \rangle$ from the data would give, with the present statistics, experimental uncertainties for these observables in the ballpark ~ 0.10 – 0.20 , as is the case for the observables in table 6.

We have also seen in section 8.2 that even with the enlarged uncertainties of table 7, the constraints on \mathcal{C}_7 , \mathcal{C}'_7 from $\langle P_2 \rangle$ are nontrivial. These uncertainties will improve very significantly in the near future in forthcoming analyses of larger data sets by the LHCb collaboration.

In order to illustrate the very large impact that moderately precise measurements of the clean observables will have on New Physics, we consider the constraints on \mathcal{C}_7 , \mathcal{C}'_7 by hypothetical measurements of $\langle P_1 \rangle_{[2,4,3]}$, $\langle P_2 \rangle_{[2,4,3]}$, $\langle P'_4 \rangle_{[2,4,3]}$ and $\langle P'_5 \rangle_{[2,4,3]}$ with central values at their SM points, and experimental uncertainties of $\sigma_{\text{exp}} = 0.10$. It is important to emphasise that these errors are not only realistic but also achievable with the current statistics. The result of this exercise is presented in figure 14, where the individual constraints in the \mathcal{C}_7 - \mathcal{C}'_7 plane (corresponding to Scenario A) from these four observables are shown together with the combined 68.3% and 95.5% C.L. regions. Clearly these observables

will play a very important role in the future, with the potentiality of providing the first unambiguous signal of New Physics in the flavour sector.

9 Comparison with other works

Several theoretical studies [6–12, 54] have exploited the recent measurements of $b \rightarrow s\gamma$ and $b \rightarrow sl^+\ell^-$. We mention here briefly some differences of recent references with respect to the present work where only low- q^2 bins were considered. With that respect, one should notice that these references use at most the $B \rightarrow K^*\ell^+\ell^-$ observables integrated over [1,6] GeV², but not the values on finer bins included here.

In ref. [54], a detailed discussion was provided concerning additional observables related to the photon polarisation in $b \rightarrow s\gamma$ transitions, allowing one to constrain \mathcal{C}_7 and \mathcal{C}'_7 more precisely. The three processes of interest were $B^0 \rightarrow K^{*0}(\rightarrow K_S\pi^0)\gamma$, $B \rightarrow K_1(\rightarrow K\pi\pi)\gamma$ and $B^0 \rightarrow K^{*0}(\rightarrow K^+\pi^-)\ell^+\ell^-$ at low q^2 , showing the potential of an analysis combining all four processes even in the case of complex Wilson coefficients (assuming that there are only small contributions from the other operators for $B^0 \rightarrow K^{*0}(\rightarrow K^+\pi^-)\ell^+\ell^-$ in this energy region). In particular, the current constraints on $\mathcal{C}_7(\prime)$ from $B(B \rightarrow X_s\gamma)$ and $S_{K^*\gamma}$ were studied, either with real or complex values for these Wilson coefficients (and SM values for the other operators). In the real case, a four-fold degeneracy is observed, corresponding essentially to the regions obtained in our Scenario A without $B \rightarrow K^*\gamma$ isospin asymmetry (see figure 4). These ambiguities (which are even more numerous in the complex-valued cases) can in principle be lifted once more observables are measured from all three processes, with different strengths and weaknesses depending on the NP scenario considered.

In ref. [12] (which updates ref. [9]), branching ratios for the inclusive modes $B \rightarrow X_s\gamma$ and $B \rightarrow X_s\ell\ell$ (low and high- q^2) and the inclusive CP asymmetry for $b \rightarrow s\gamma$ are considered, as well as the branching ratio for $B \rightarrow K\ell\ell$ and $B_s \rightarrow \mu^+\mu^-$ and observables for $B \rightarrow K^*\mu^+\mu^-$ (branching ratio, longitudinal polarization fraction, forward-backward asymmetry, A_9 and S_3 both in low- and high- q^2 regions). The constraints are put on $\mathcal{C}_7(\prime), \mathcal{C}_9(\prime), \mathcal{C}_{10}(\prime), \mathcal{C}_S(\prime), \mathcal{C}_P(\prime)$ with real or complex values, first in a frequentist framework, then in a Bayesian approach, with a focus on some specific NP scenarios. Once again, a good agreement with the SM is obtained. In addition, the possibility of a sign-flip in \mathcal{C}_7 alone is disfavoured due to the branching ratio of $B \rightarrow X_s\gamma$ as well as A_{FB} . In the scenario without right-handed currents and with real values of the coefficients of the SM operators, the flipped sign solution $\mathcal{C}_{7,9,10} \rightarrow -\mathcal{C}_{7,9,10}$ remains allowed (since the $B \rightarrow K^*\gamma$ isospin asymmetry was not included in the analysis). In this scenario, the constraints obtained on the Wilson coefficients are similar to those that we obtain for scenario B, up to additional solutions with $\mathcal{C}'_7 \neq 0$. The very large parameter space and the choice of different scenarios prevents us from comparing our results in more detail, but we stress that the analysis in ref. [12] includes form-factor sensitive observables like S_3 .

In ref. [11], the authors did not consider the inclusive modes $B \rightarrow X_s\gamma$ and $B \rightarrow X_s\ell\ell$. On the other hand, they considered the branching ratio of the exclusive modes $B \rightarrow K^*\gamma$, $B \rightarrow K\ell\ell$, $B_s \rightarrow \mu^+\mu^-$ as well as the available observables for $B \rightarrow K^*\mu^+\mu^-$ (branching

ratio, longitudinal fraction, forward-backward asymmetry, transverse asymmetry $A_T^{(2)}$ and S_3 both in low- and high- q^2 regions). The constraints were analysed in the SM basis, constraining only real values of $\mathcal{C}_7, \mathcal{C}_9, \mathcal{C}_{10}$, in a Bayesian framework. The inputs for the hadronic form factors are taken from the light-cone sum rule analysis of ref. [28]. They provide also predictions for the other transverse asymmetries $A_T^{(i)}$, as well as for the low-recoil observables $H_T^{(i)}$ introduced in ref. [35]. Obviously, as the list of inputs, scenarios and statistical frameworks are rather different, we can only perform a limited comparison with our benchmark points in different NP scenarios. In both analyses, the SM point is favoured. A second solution, with flipped signs $\mathcal{C}_{7,9,10} \rightarrow -\mathcal{C}_{7,9,10}$ is allowed in ref. [11] as in the previous references. The authors observe an update from prior to posterior p.d.f.'s of the form factors, which can be interpreted as the fact that the data themselves constrain partly the hadronic uncertainties, due to a slight tension between $B \rightarrow K\ell\ell$ and $B \rightarrow K^*\ell\ell$. In view of our discussion concerning theoretically clean and form-factor sensitive observables, it would be interesting to perform a similar analysis separating the two sets of observables.

10 Conclusions

The decay mode $B \rightarrow K^*\ell^+\ell^-$, with its large set of angular observables, is becoming a more and more important process in constraining New Physics models. These constraints are complementary to those from inclusive and exclusive radiative decays. With increasingly precise experimental data on these modes and the prospects for the future from the LHC, theoretical uncertainties must be kept under control. In this context, the use of theoretically clean observables is not only convenient but also extremely recommendable, or even mandatory. The sensitivity of different observables to hadronic uncertainties has been addressed in section 6, and the conclusion is that clean observables such as P_1 are far more robust than other observables like F_L or S_3 , translating into a better performance in discriminating among different models.

A full description of the angular distribution of the $B \rightarrow K^*\ell^+\ell^-$ decay in terms of a maximum set of clean observables was presented in ref. [32], where the observables $P_{1,2,3,4,5,6}$ were introduced. In this paper we have presented a simple and compact expression for the coefficients of the distribution in terms of these clean observables, and given SM predictions for these observables integrated over a series of q^2 -bins, that can be directly compared with experimental data. These predictions are collected in table 1 and in figures 1 and 2. As an important point, we have seen that the first three clean observables $P_{1,2,3}$ are *already measured*, and can be extracted from the latest measurements by the LHCb collaboration [4]. The experimental numbers for these observables together with their SM predictions are given in table 7. The uncertainties attached to these numbers are certainly overestimated, as we did not have the required correlations among experimental measurements (and thus treated them as uncorrelated).

We have also studied the model-independent constraints on the Wilson coefficients $\mathcal{C}_7^{(\prime)}$, $\mathcal{C}_9^{(\prime)}$, $\mathcal{C}_{10}^{(\prime)}$ from radiative decays $B \rightarrow X_s\gamma$ and $B \rightarrow K^*\gamma$, and semileptonic decays $B \rightarrow X_s\mu^+\mu^-$ and $B \rightarrow K^*\mu^+\mu^-$. Excluding all $B \rightarrow K^*\mu^+\mu^-$ observables except for the integrated observables $\langle A_{FB} \rangle_{[1,6]}$ and $\langle F_L \rangle_{[1,6]}$ leads to the set of constraints shown

in figures 4–7. Identifying a set of benchmark New Physics points compatible with those bounds, one can see that very large New Physics contributions to other observables in $B \rightarrow K^* \mu^+ \mu^-$ are allowed, specially in P_1 , P_2 , P_4' and P_5' (see figure 9). We have investigated the present constraints imposed by P_1 , P_2 and P_3 on C_7 , C_7' , with already quite interesting constraints from P_2 suggesting a lower value of C_7 than the SM value. We then showed the powerful prospects that the set of clean observables $P_{1,2,3}$ and $P_{4',5',6'}$ will have in the short term to discriminate possible New Physics contributions, illustrated in figure 14.

Considering the advantages provided by the use of the P_i observables at large recoil, we hope that the present study will be a strong incentive for experimentalists to rephrase their study of the low- q^2 region in terms of these observables.

Acknowledgments

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A J_8 in terms of P_i observables and the Q observable

In section 2 we have provided the explicit expressions for the coefficients of the distribution in terms of the observables of the basis, neglecting mass terms and scalar contributions.⁶ In this case, one finds very simple and compact expressions for all these coefficients (see eqs. (2.13)) with the exception of J_8 . The reason is that J_8 , in the absence of scalar contributions, is *not* an independent quantity (exactly like $J_{1s} = 3J_{2s}$ and $J_{1c} = -J_{2c}$ in the massless case) and deserves a separate discussion. The counting of degrees of freedom and continuous symmetries in this case shows that there are only 8 degrees of freedom parametrised by the observables $P_{i=1\dots6}$, F_L and $d\Gamma/dq^2$. This means that J_8 can be expressed in terms of these observables:

$$J_8 = -\sqrt{\frac{F_T F_L}{1 - P_1}} \frac{d\Gamma}{dq^2} \left\{ (P_2 P_6 - P_3 P_4) + \eta \left((P_2 P_6 - P_3 P_4)^2 + P_5 (P_2 P_4 + P_3 P_6) \sqrt{1 - P_1^2} \right. \right. \\ \left. \left. + \frac{1}{4} \left(1 - \sum_{i=4}^6 P_i^2 \right) (1 - P_1^2) - P_2^2 - P_3^2 \right)^{\frac{1}{2}} \right\} \quad (\text{A.1})$$

This expression is found by solving J_8 in terms of the other coefficients using the relation in eq. (3.15) of ref. [33] together with eqs. (2.13). One can also replace $P_{4,5,6}$ by $P'_{4,5,6}$ using eqs. (2.14)–(2.16).

We notice that in eq. (A.1), a discrete quantity η is left as a free parameter that can take only two values ± 1 in the massless case. This parameter is indeed an observable, and its SM prediction is $\eta^{\text{SM}} = +1$ as can be seen by substituting the SM values for the observables P_i . This has the interesting consequence that a measurement of $\eta = -1$ would

⁶The general case including lepton masses and scalars is discussed in full detail in ref. [32].

be an unambiguous indication of New Physics (originating, for instance, from new weak phases or sign flips in Wilson coefficients). Deviations from $|\eta| = 1$ could be expected, for instance, from scalar contributions entering P_5 and P_6 , since such contributions would break the symmetry relation among the coefficients of the angular distribution.

However, when one tries to write a similar relationship for “binned” observables, it is clear that a naive substitution $P_i \rightarrow \langle P_i \rangle_{\text{bin}}$ is not possible due to the highly non-linear form of eq. (A.1). One practical solution is to introduce an extra clean observable Q (or Q'):

$$J_8 = -\frac{1}{2}Q' \sqrt{F_T F_L} \frac{d\Gamma}{dq^2}, \quad (\text{A.2})$$

with $Q' = Q\sqrt{1+P_1}$ and whose definition, also valid in the massive case, is

$$Q = \frac{\text{Im}(n_0^\dagger n_\perp)}{\sqrt{|n_0|^2 |n_\perp|^2}} = -\frac{\sqrt{2}J_8}{\sqrt{-J_{2c}(2J_{2s} + J_3)}}. \quad (\text{A.3})$$

The vectors n_i (with $i = 0, \perp$) are defined in ref. [32]. This clean observable Q (or Q') is related to the form factor-sensitive observable S_8 by means of

$$Q' = -\frac{2S_8}{\sqrt{F_T F_L}}. \quad (\text{A.4})$$

For completeness we also provide the expression of this observable in terms of our basis of observables:

$$Q = \frac{2}{\sqrt{1-P_1^2}} \left\{ (P_2 P_6 - P_3 P_4) + \eta \left((P_2 P_6 - P_3 P_4)^2 + P_5 (P_2 P_4 + P_3 P_6) \sqrt{1-P_1^2} \right. \right. \\ \left. \left. + \frac{1}{4} \left(1 - \sum_{i=4}^6 P_i^2 \right) (1 - P_1^2) - P_2^2 - P_3^2 \right)^{\frac{1}{2}} \right\}. \quad (\text{A.5})$$

In principle, this implies that Q (or Q') can be measured using either eq. (A.3) or eq. (A.5) (with $\eta = +1$ in the SM but free in general). However, eq. (A.5) would be of practical experimental use only in the limit of the size of the binning going to zero. Of course, eq. (A.5) can be very easily turned into an equation that relates S_8 with the other S_i by using eq. (2.4) and eq. (2.17). This shows the redundancy of S_8 in the massless case, again up to a single discrete parameter η .

At a more practical level, the integrated form of this observable from eq. (A.3) is:

$$\langle Q' \rangle_{\text{bin}} = \frac{-2 \int_{\text{bin}} dq^2 J_8(q^2)}{\sqrt{\int_{\text{bin}} dq^2 c_4(q^2) \int_{\text{bin}} dq^2 (c_0(q^2) - c_4(q^2))}} = \frac{-2 \langle J_8 \rangle_{\text{bin}}}{\sqrt{\langle c_4 \rangle_{\text{bin}} (\langle c_0 \rangle_{\text{bin}} - \langle c_4 \rangle_{\text{bin}})}} \quad (\text{A.6})$$

and its SM prediction is given in figure 15. Notice that, as mentioned above, the binning procedure breaks the relation eq. (3.15) of ref. [33] among the coefficients of the distribution, which is only recovered in the limit of the size of the binning going to zero, when the observable η would become accessible.

In conclusion the set of eqs. (2.13) and eq. (A.1) provides a complete parametrisation of the distribution in the massless case with no scalars in terms of only eight parameters (eight

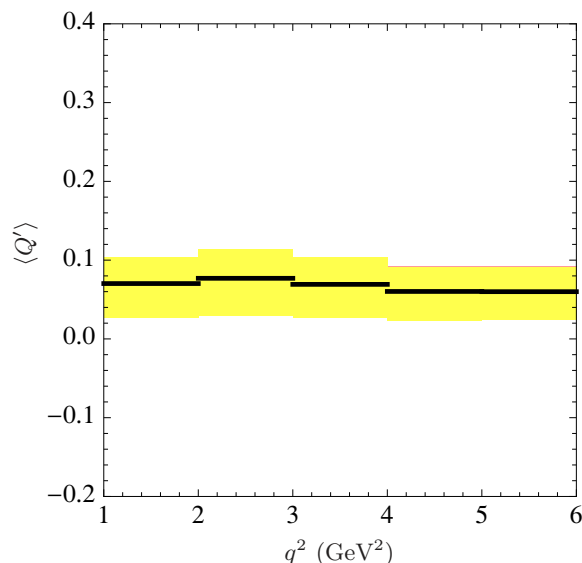


Figure 15. SM prediction for the observable Q' , integrated in 5 bins of width 1 GeV².

observables $P_{1,2,3,4,5,6}$ and $F_L, d\Gamma/dq^2$) in agreement with the eight independent degrees of freedom, but there is one extra redundant Q (or Q') observable (up to a discrete parameter η) which however can be fixed only once $\langle J_8 \rangle_{\text{bin}}$ is measured, due to the binning procedure adopted by experimental analyses.

B New Physics expressions for binned observables

In this appendix we present the numerical expressions for the integrated observables $\langle A_{\text{FB}} \rangle$, $\langle F_L \rangle$, $\langle P_{1,2,3} \rangle$ and $\langle P'_{4,5,6} \rangle$ as a function of the NP Wilson coefficients, for different choices of the q^2 -binning.

The formula for any observable $\langle \mathcal{O}_k \rangle$ has the general form:

$$\langle \mathcal{O}_k \rangle = \frac{\sum_{i,j=0,7,7',9,9',10,10'} N_{(i,j)} \delta \mathcal{C}_i \delta \mathcal{C}_j}{\sum_{i,j=0,7,7',9,9',10,10'} D_{(i,j)} \delta \mathcal{C}_i \delta \mathcal{C}_j} \pm \delta_k \quad (\text{B.1})$$

where $\delta C_0 \equiv 1$. The coefficients δ , $N_{i,j}$ and $D_{i,j}$ corresponding to each observable are collected in tables 8–12, where only nonzero coefficients are displayed. The coefficient $\delta \mathcal{C}_i$ denotes the NP contribution to the Wilson coefficient \mathcal{C}_i at the hadronic scale μ_b . The parameters δ_k correspond to the theoretical error assigned to each observable, and are also collected in tables 8–12. We assume here that the uncertainty obtained within the SM can be considered as a good estimate of the uncertainty for arbitrary (NP) values of the Wilson coefficients. As discussed in detail in section 6, this assumption is appropriate for theoretically clean observables only.

$\langle A_{FB} \rangle$								
	[1, 6]	[2, 4, 3]	[4.3, 8.68]	[1, 2]	[2, 3]	[3, 4]	[4, 5]	[5, 6]
δ	0.035	0.059	0.126	0.127	0.106	0.04	0.051	0.089
$N_{(0,0)}$	-132.63	-129.81	1075.21	-169.19	-103.04	-31.81	44.25	127.17
$N_{(0,7)}$	3659.89	1683.77	3248.78	709.73	725.48	736.	742.78	745.9
$N_{(0,7')}$	53.45	24.59	47.45	10.37	10.6	10.75	10.85	10.89
$N_{(0,9)}$	273.22	112.45	446.79	22.53	38.38	54.54	70.81	86.96
$N_{(0,10)}$	30.78	30.13	-249.55	39.27	23.92	7.38	-10.27	-29.51
$N_{(0,10')}$	-9.55	-4.44	-7.25	-1.92	-1.93	-1.93	-1.91	-1.85
$N_{(7,7')}$	-1210.54	-544.54	-1139.94	-242.44	-234.51	-237.58	-244.07	-251.93
$N_{(7,9)}$	1502.03	688.89	1366.	286.75	295.32	301.92	307.12	310.92
$N_{(7,9')}$	-415.72	-187.6	-424.87	-73.44	-78.45	-83.25	-87.97	-92.62
$N_{(7',9)}$	-448.14	-202.46	-454.37	-79.62	-84.82	-89.76	-94.6	-99.33
$N_{(7',9')}$	1511.64	693.23	1375.74	288.45	297.14	303.84	309.14	313.05
$N_{(9,9')}$	-182.43	-82.42	-184.94	-32.42	-34.53	-36.54	-38.5	-40.43
$N_{(10,10')}$	-182.43	-82.42	-184.94	-32.42	-34.53	-36.54	-38.5	-40.43
$D_{(0,0)}$	3930.05	1652.9	4881.21	788.91	693.77	727.04	805.92	914.41
$D_{(0,7)}$	-675.8	-133.39	3759.72	-1510.4	-367.15	131.35	427.89	642.5
$D_{(0,7')}$	-1514.33	-676.99	-1832.12	-243.59	-277.12	-303.71	-329.95	-359.96
$D_{(0,9)}$	607.74	251.22	1056.34	62.61	90.17	119.3	150.44	185.22
$D_{(0,9')}$	-659.	-294.16	-766.64	-114.1	-122.32	-130.82	-140.19	-151.57
$D_{(0,10)}$	-1121.92	-494.21	-1333.74	-167.85	-196.49	-224.8	-252.76	-280.02
$D_{(0,10')}$	785.99	355.1	796.84	139.7	148.79	157.42	165.89	174.19
$D_{(7,7)}$	19751.1	8522.58	8359.39	7326.88	4445.59	3246.87	2580.52	2151.26
$D_{(7,7')}$	19812.1	8549.86	8412.14	7340.65	4457.7	3258.55	2592.18	2163.05
$D_{(9,9)}$	260.39	114.7	309.56	38.96	45.6	52.18	58.66	64.99
$D_{(9',9')}$	260.39	114.7	309.56	38.96	45.6	52.18	58.66	64.99
$D_{(10,10)}$	260.39	114.7	309.56	38.96	45.6	52.18	58.66	64.99
$D_{(10',10')}$	260.39	114.7	309.56	38.96	45.6	52.18	58.66	64.99
$D_{(7,7')}$	-1210.54	-544.54	-1139.94	-242.44	-234.51	-237.58	-244.07	-251.93
$D_{(7,9)}$	1502.03	688.89	1366.	286.75	295.32	301.92	307.12	310.92
$D_{(7,9')}$	-415.72	-187.6	-424.87	-73.44	-78.45	-83.25	-87.97	-92.62
$D_{(7',9)}$	-448.14	-202.46	-454.37	-79.62	-84.82	-89.76	-94.6	-99.33
$D_{(7',9')}$	1511.64	693.23	1375.74	288.45	297.14	303.84	309.14	313.05
$D_{(9,9')}$	-182.43	-82.42	-184.94	-32.42	-34.53	-36.54	-38.5	-40.43
$D_{(10,10')}$	-182.43	-82.42	-184.94	-32.42	-34.53	-36.54	-38.5	-40.43

Table 8. Coefficients for the New Physics formula of $\langle A_{FB} \rangle$.

$\langle F_L \rangle$								
δ	[1, 6]	[2, 4.3]	[4.3, 8.68]	[1, 2]	[2, 3]	[3, 4]	[4, 5]	[5, 6]
δ	0.178	0.162	0.18	0.211	0.164	0.158	0.169	0.179
$N_{(0,0)}$	1413.04	635.91	1559.39	251.77	266.81	281.63	297.22	315.6
$N_{(0,7)}$	681.08	305.96	758.66	120.71	128.16	135.6	143.55	153.06
$N_{(0,7')}$	-696.01	-312.66	-775.28	-123.35	-130.97	-138.58	-146.69	-156.42
$N_{(0,9)}$	323.9	145.53	360.32	57.46	60.98	64.49	68.24	72.74
$N_{(0,9')}$	-323.9	-145.53	-360.32	-57.46	-60.98	-64.49	-68.24	-72.74
$N_{(0,10)}$	-378.21	-171.92	-367.37	-68.88	-72.6	-75.94	-79.01	-81.79
$N_{(0,10')}$	378.21	171.92	367.37	68.88	72.6	75.94	79.01	81.79
$N_{(7,7)}$	388.11	176.38	377.96	70.55	74.44	77.92	81.14	84.06
$N_{(7',7')}$	405.3	184.19	394.71	73.68	77.73	81.38	84.73	87.79
$N_{(9,9)}$	87.78	39.9	85.27	15.99	16.85	17.62	18.34	18.98
$N_{(9',9')}$	87.78	39.9	85.27	15.99	16.85	17.62	18.34	18.98
$N_{(10,10)}$	87.78	39.9	85.27	15.99	16.85	17.62	18.34	18.98
$N_{(10',10')}$	87.78	39.9	85.27	15.99	16.85	17.62	18.34	18.98
$N_{(7,7')}$	-605.27	-272.27	-569.97	-121.22	-117.25	-118.79	-122.04	-125.97
$N_{(7,9)}$	751.01	344.45	683.	143.37	147.66	150.96	153.56	155.46
$N_{(7,9')}$	-207.86	-93.8	-212.44	-36.72	-39.23	-41.62	-43.98	-46.31
$N_{(7',9)}$	-224.07	-101.23	-227.18	-39.81	-42.41	-44.88	-47.3	-49.67
$N_{(7',9')}$	755.82	346.62	687.87	144.23	148.57	151.92	154.57	156.53
$N_{(9,9')}$	-91.21	-41.21	-92.47	-16.21	-17.27	-18.27	-19.25	-20.21
$N_{(10,10')}$	-91.21	-41.21	-92.47	-16.21	-17.27	-18.27	-19.25	-20.21
$D_{(0,0)}$	1965.02	826.45	2440.6	394.45	346.88	363.52	402.96	457.21
$D_{(0,7)}$	-337.9	-66.69	1879.86	-755.2	-183.57	65.68	213.94	321.25
$D_{(0,7')}$	-757.17	-338.49	-916.06	-121.8	-138.56	-151.85	-164.98	-179.98
$D_{(0,9)}$	303.87	125.61	528.17	31.31	45.09	59.65	75.22	92.61
$D_{(0,9')}$	-329.5	-147.08	-383.32	-57.05	-61.16	-65.41	-70.1	-75.79
$D_{(0,10)}$	-560.96	-247.1	-666.87	-83.93	-98.25	-112.4	-126.38	-140.01
$D_{(0,10')}$	393.	177.55	398.42	69.85	74.4	78.71	82.94	87.09
$D_{(7,7)}$	9875.55	4261.29	4179.7	3663.44	2222.8	1623.43	1290.26	1075.63
$D_{(7',7')}$	9906.06	4274.93	4206.07	3670.32	2228.85	1629.28	1296.09	1081.53
$D_{(9,9)}$	130.2	57.35	154.78	19.48	22.8	26.09	29.33	32.5
$D_{(9',9')}$	130.2	57.35	154.78	19.48	22.8	26.09	29.33	32.5
$D_{(10,10)}$	130.2	57.35	154.78	19.48	22.8	26.09	29.33	32.5
$D_{(10',10')}$	130.2	57.35	154.78	19.48	22.8	26.09	29.33	32.5
$D_{(7,7')}$	-605.27	-272.27	-569.97	-121.22	-117.25	-118.79	-122.04	-125.97
$D_{(7,9)}$	751.01	344.45	683.	143.37	147.66	150.96	153.56	155.46
$D_{(7,9')}$	-207.86	-93.8	-212.44	-36.72	-39.23	-41.62	-43.98	-46.31
$D_{(7',9)}$	-224.07	-101.23	-227.18	-39.81	-42.41	-44.88	-47.3	-49.67
$D_{(7',9')}$	755.82	346.62	687.87	144.23	148.57	151.92	154.57	156.53
$D_{(9,9')}$	-91.21	-41.21	-92.47	-16.21	-17.27	-18.27	-19.25	-20.21
$D_{(10,10')}$	-91.21	-41.21	-92.47	-16.21	-17.27	-18.27	-19.25	-20.21

Table 9. Coefficients for the New Physics formula of $\langle F_L \rangle$.

$\langle P_1 \rangle$								
	[1, 6]	[2, 4.3]	[4.3, 8.68]	[1, 2]	[2, 3]	[3, 4]	[4, 5]	[5, 6]
δ	0.051	0.05	0.06	0.052	0.056	0.046	0.046	0.053
$N_{(0,0)}$	-30.44	-9.75	-103.1	1.11	-1.96	-5.4	-9.53	-14.65
$N_{(0,7)}$	-82.98	-33.81	-116.67	-17.23	-14.27	-14.77	-16.77	-19.95
$N_{(0,7')}$	-1017.67	-372.1	1124.22	-875.94	-311.58	-69.64	70.79	168.7
$N_{(0,9)}$	-5.6	-1.55	-23.	0.41	-0.18	-0.92	-1.85	-3.05
$N_{(0,9')}$	-20.03	-19.92	167.85	-26.15	-15.89	-4.84	6.98	19.87
$N_{(0,10)}$	14.79	5.63	31.05	0.98	1.8	2.78	3.94	5.3
$N_{(0,10')}$	-182.75	-75.18	-299.5	-15.05	-25.65	-36.46	-47.37	-58.22
$N_{(7,7)}$	-5.14	-4.4	-102.13	27.93	4.9	-6.01	-13.18	-18.78
$N_{(7',7')}$	-412.36	-179.74	-265.45	-126.24	-87.31	-72.35	-65.09	-61.37
$N_{(9,9)}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$N_{(9',9')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$N_{(10,10)}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$N_{(10',10')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$N_{(7,7')}$	-208.65	-92.03	-183.73	-49.13	-41.19	-39.16	-39.12	-40.06
$N_{(7,9)}$	566.44	260.55	503.48	109.79	112.25	113.9	114.99	115.51
$N_{(7,9')}$	-23.29	-9.91	-32.92	-3.14	-3.81	-4.56	-5.41	-6.36
$N_{(7',9)}$	-35.45	-15.5	-43.73	-5.49	-6.22	-7.01	-7.88	-8.84
$N_{(7',9')}$	567.2	260.89	504.42	109.91	112.38	114.05	115.15	115.7
$N_{(9,9')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$N_{(10,10')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$D_{(0,0)}$	551.99	190.54	881.21	142.68	80.07	81.89	105.74	141.61
$D_{(0,7)}$	-1018.98	-372.65	1121.2	-875.91	-311.74	-69.93	70.4	168.19
$D_{(0,7')}$	-61.16	-25.83	-140.78	1.56	-7.59	-13.28	-18.28	-23.56
$D_{(0,9)}$	-20.03	-19.92	167.85	-26.15	-15.89	-4.84	6.98	19.87
$D_{(0,9')}$	-5.6	-1.55	-23.	0.41	-0.18	-0.92	-1.85	-3.05
$D_{(0,10)}$	-182.75	-75.18	-299.5	-15.05	-25.65	-36.46	-47.37	-58.22
$D_{(0,10')}$	14.79	5.63	31.05	0.98	1.8	2.78	3.94	5.3
$D_{(7,7)}$	9487.44	4084.91	3801.73	3592.89	2148.36	1545.51	1209.12	991.57
$D_{(7',7')}$	9500.76	4090.74	3811.36	3596.65	2151.12	1547.9	1211.35	993.74
$D_{(9,9)}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(9',9')}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(10,10)}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(10',10')}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(7,7')}$	-208.65	-92.03	-183.73	-49.13	-41.19	-39.16	-39.12	-40.06
$D_{(7,9)}$	566.44	260.55	503.48	109.79	112.25	113.9	114.99	115.51
$D_{(7,9')}$	-23.29	-9.91	-32.92	-3.14	-3.81	-4.56	-5.41	-6.36
$D_{(7',9)}$	-35.45	-15.5	-43.73	-5.49	-6.22	-7.01	-7.88	-8.84
$D_{(7',9')}$	567.2	260.89	504.42	109.91	112.38	114.05	115.15	115.7
$D_{(9,9')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$D_{(10,10')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23

Table 10. Coefficients for the New Physics formula of $\langle P_1 \rangle$.

$\langle P_2 \rangle$								
	[1, 6]	[2, 4, 3]	[4.3, 8.68]	[1, 2]	[2, 3]	[3, 4]	[4, 5]	[5, 6]
δ	0.067	0.07	0.074	0.023	0.032	0.085	0.082	0.071
$N_{(0,0)}$	44.21	43.27	-358.4	56.4	34.35	10.6	-14.75	-42.39
$N_{(0,7)}$	-1219.96	-561.26	-1082.93	-236.58	-241.83	-245.33	-247.59	-248.63
$N_{(0,7')}$	-17.82	-8.2	-15.82	-3.46	-3.53	-3.58	-3.62	-3.63
$N_{(0,9)}$	-91.07	-37.48	-148.93	-7.51	-12.79	-18.18	-23.6	-28.99
$N_{(0,10)}$	-10.26	-10.04	83.18	-13.09	-7.97	-2.46	3.42	9.84
$N_{(0,10')}$	3.18	1.48	2.42	0.64	0.64	0.64	0.64	0.62
$N_{(7,7')}$	-208.65	-92.03	-183.73	-49.13	-41.19	-39.16	-39.12	-40.06
$N_{(7,9)}$	566.44	260.55	503.48	109.79	112.25	113.9	114.99	115.51
$N_{(7,9')}$	-23.29	-9.91	-32.92	-3.14	-3.81	-4.56	-5.41	-6.36
$N_{(7',9)}$	-35.45	-15.5	-43.73	-5.49	-6.22	-7.01	-7.88	-8.84
$N_{(7',9')}$	567.2	260.89	504.42	109.91	112.38	114.05	115.15	115.7
$N_{(9,9')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$N_{(10,10')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$D_{(0,0)}$	551.99	190.54	881.21	142.68	80.07	81.89	105.74	141.61
$D_{(0,7)}$	-1018.98	-372.65	1121.2	-875.91	-311.74	-69.93	70.4	168.19
$D_{(0,7')}$	-61.16	-25.83	-140.78	1.56	-7.59	-13.28	-18.28	-23.56
$D_{(0,9)}$	-20.03	-19.92	167.85	-26.15	-15.89	-4.84	6.98	19.87
$D_{(0,9')}$	-5.6	-1.55	-23.	0.41	-0.18	-0.92	-1.85	-3.05
$D_{(0,10)}$	-182.75	-75.18	-299.5	-15.05	-25.65	-36.46	-47.37	-58.22
$D_{(0,10')}$	14.79	5.63	31.05	0.98	1.8	2.78	3.94	5.3
$D_{(7,7)}$	9487.44	4084.91	3801.73	3592.89	2148.36	1545.51	1209.12	991.57
$D_{(7',7')}$	9500.76	4090.74	3811.36	3596.65	2151.12	1547.9	1211.35	993.74
$D_{(9,9)}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(9',9')}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(10,10)}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(10',10')}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(7,7')}$	-208.65	-92.03	-183.73	-49.13	-41.19	-39.16	-39.12	-40.06
$D_{(7,9)}$	566.44	260.55	503.48	109.79	112.25	113.9	114.99	115.51
$D_{(7,9')}$	-23.29	-9.91	-32.92	-3.14	-3.81	-4.56	-5.41	-6.36
$D_{(7',9)}$	-35.45	-15.5	-43.73	-5.49	-6.22	-7.01	-7.88	-8.84
$D_{(7',9')}$	567.2	260.89	504.42	109.91	112.38	114.05	115.15	115.7
$D_{(9,9')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$D_{(10,10')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23

Table 11. Coefficients for the New Physics formula of $\langle P_2 \rangle$.

$\langle P_3 \rangle$								
	[1, 6]	[2, 4.3]	[4.3, 8.68]	[1, 2]	[2, 3]	[3, 4]	[4, 5]	[5, 6]
δ	0.024	0.024	0.027	0.027	0.027	0.022	0.022	0.024
$N_{(0,0)}$	-1.73	-0.72	-0.57	-0.49	-0.34	-0.3	-0.28	-0.32
$N_{(0,7)}$	3.7	1.54	0.98	1.29	0.79	0.59	0.51	0.52
$N_{(0,7')}$	172.63	71.81	45.75	60.34	36.74	27.69	23.8	24.07
$N_{(0,9)}$	-0.16	-0.07	-0.07	-0.03	-0.03	-0.03	-0.03	-0.04
$N_{(0,9')}$	10.89	4.61	4.74	1.84	1.92	2.04	2.27	2.81
$N_{(7,7')}$	-208.65	-92.03	-183.73	-49.13	-41.19	-39.16	-39.12	-40.06
$N_{(7,9)}$	566.44	260.55	503.48	109.79	112.25	113.9	114.99	115.51
$N_{(7,9')}$	-23.29	-9.91	-32.92	-3.14	-3.81	-4.56	-5.41	-6.36
$N_{(7',9)}$	-35.45	-15.5	-43.73	-5.49	-6.22	-7.01	-7.88	-8.84
$N_{(7',9')}$	567.2	260.89	504.42	109.91	112.38	114.05	115.15	115.7
$N_{(9,9')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$N_{(10,10')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$D_{(0,0)}$	551.99	190.54	881.21	142.68	80.07	81.89	105.74	141.61
$D_{(0,7)}$	-1018.98	-372.65	1121.2	-875.91	-311.74	-69.93	70.4	168.19
$D_{(0,7')}$	-61.16	-25.83	-140.78	1.56	-7.59	-13.28	-18.28	-23.56
$D_{(0,9)}$	-20.03	-19.92	167.85	-26.15	-15.89	-4.84	6.98	19.87
$D_{(0,9')}$	-5.6	-1.55	-23.	0.41	-0.18	-0.92	-1.85	-3.05
$D_{(0,10)}$	-182.75	-75.18	-299.5	-15.05	-25.65	-36.46	-47.37	-58.22
$D_{(0,10')}$	14.79	5.63	31.05	0.98	1.8	2.78	3.94	5.3
$D_{(7,7)}$	9487.44	4084.91	3801.73	3592.89	2148.36	1545.51	1209.12	991.57
$D_{(7',7')}$	9500.76	4090.74	3811.36	3596.65	2151.12	1547.9	1211.35	993.74
$D_{(9,9)}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(9',9')}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(10,10)}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(10',10')}$	42.42	17.45	69.51	3.49	5.95	8.46	10.99	13.51
$D_{(7,7')}$	-208.65	-92.03	-183.73	-49.13	-41.19	-39.16	-39.12	-40.06
$D_{(7,9)}$	566.44	260.55	503.48	109.79	112.25	113.9	114.99	115.51
$D_{(7,9')}$	-23.29	-9.91	-32.92	-3.14	-3.81	-4.56	-5.41	-6.36
$D_{(7',9)}$	-35.45	-15.5	-43.73	-5.49	-6.22	-7.01	-7.88	-8.84
$D_{(7',9')}$	567.2	260.89	504.42	109.91	112.38	114.05	115.15	115.7
$D_{(9,9')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23
$D_{(10,10')}$	-3.43	-1.31	-7.21	-0.23	-0.42	-0.64	-0.91	-1.23

Table 12. Coefficients for the New Physics formula of $\langle P_3 \rangle$.

C Statistical approach

We determine and combine our constraints in a frequentist framework, treating theoretical and experimental uncertainties on the same footing (i.e., taking them as normally distributed random variables). The model independent bounds are obtained in the following way. A chi-square function is constructed, according to:

$$\chi^2(p) = \sum_k \frac{(\mathcal{O}_k^{th}(p) - \mathcal{O}_k^{\text{exp}})^2}{\delta_k^2 + \sigma_k^2} \quad (\text{C.1})$$

where p are the theoretical parameters constrained by the analysis (Wilson coefficients, hadronic quantities, CKM matrix elements. . .), \mathcal{O}_k^{th} , δ_k are the central value for theoretical prediction and error for the observable \mathcal{O}_k , and $\mathcal{O}_k^{\text{exp}} \pm \sigma_k$ is the experimental average. In case of several uncertainties (statistical and systematic ones, for instance), we combine them in quadrature. In the case of the observables $\langle A_{\text{FB}} \rangle$, $\langle F_L \rangle$, $\langle P_{1,2,3} \rangle$ and $\langle P'_{4,5,6} \rangle$, the numbers \mathcal{O}_k^{th} and δ_k are given by eq. (B.1). We add to the χ^2 similar quadratic terms for the theoretical quantities involved (decay constants, form factors, quark masses, CKM matrix elements).

We want to obtain the constraints in the two-dimensional plane corresponding to two (shifts of) Wilson coefficients $(\delta\mathcal{C}_a, \delta\mathcal{C}_b)$ among all the parameters p (we denote the remaining theoretical parameters q). We obtain these regions by drawing the region of $(\delta\mathcal{C}_a, \delta\mathcal{C}_b)$ where:

$$\min_q \chi^2(\delta\mathcal{C}_a, \delta\mathcal{C}_b) - \min_p \chi^2 < \delta \quad (\text{C.2})$$

where δ depends on the dimension of the parameter space where the region is drawn (here, a two-dimensional plane), and the required confidence level (here, 2.3 for 68.3% C.L. and 6.18 for 95.5% C.L., see the review of statistics and Monte Carlo techniques in ref. [46])

At the practical level, one could in principle compute the χ^2 on a grid for the Wilson coefficients, performing the minimisation over all the nuisance parameters. One would obtain the two-dimensional contours for a given pair of Wilson coefficients $(\delta\mathcal{C}_a, \delta\mathcal{C}_b)$ by performing a further minimisation on the other Wilson coefficients. It turns out that one can improve the accuracy of the method by sampling the parameter space through a Metropolis-Hastings Markov-Chain Monte Carlo algorithm with a weight that favours the minima of the χ^2 , for instance $\exp(-\chi^2/2)$. Once this sampling has been performed, one creates the grid of points (and compute the value of the χ^2 points) by considering all the points sampled by the algorithm. The points of grid which have not been sampled at all are given a very large χ^2 . One then proceeds to a smearing procedure to get a smooth dependence of the reconstructed χ^2 grid.

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References

- [1] BELLE collaboration, J.-T. Wei et al., *Measurement of the differential branching fraction and forward-backward asymmetry for $B \rightarrow K^{(*)}\ell^+\ell^-$* , *Phys. Rev. Lett.* **103** (2009) 171801 [[arXiv:0904.0770](#)] [[INSPIRE](#)].
- [2] CDF collaboration, T. Aaltonen et al., *Measurements of the angular distributions in the decays $B \rightarrow K^{(*)}\mu^+\mu^-$ at CDF*, *Phys. Rev. Lett.* **108** (2012) 081807 [[arXiv:1108.0695](#)] [[INSPIRE](#)].
- [3] BABAR collaboration, J. Lees et al., *Measurement of branching fractions and rate asymmetries in the rare decays $B \rightarrow K^{(*)}\ell^+\ell^-$* , *Phys. Rev. D* **86** (2012) 032012 [[arXiv:1204.3933](#)] [[INSPIRE](#)].
- [4] LHCb collaboration, *Differential branching fraction and angular analysis of the $B^0 \rightarrow K^{*0}\mu^+\mu^-$ decay*, LHCb-CONF-2012-008 (2012).
- [5] LHCb collaboration, R. Aaij et al., *Measurement of the isospin asymmetry in $B \rightarrow K^{(*)}\mu^+\mu^-$ decays*, *JHEP* **07** (2012) 133 [[arXiv:1205.3422](#)] [[INSPIRE](#)].
- [6] S. Descotes-Genon, D. Ghosh, J. Matias and M. Ramon, *Exploring new physics in the C_7-C_7' plane*, *JHEP* **06** (2011) 099 [[arXiv:1104.3342](#)] [[INSPIRE](#)].
- [7] S. Descotes-Genon, D. Ghosh, J. Matias and M. Ramon, *Exploring new physics in C_7-C_7'* , *PoS(EPS-HEP2011)170* [[arXiv:1202.2172](#)] [[INSPIRE](#)].
- [8] C. Bobeth, G. Hiller and D. van Dyk, *More benefits of semileptonic rare B decays at low recoil: CP -violation*, *JHEP* **07** (2011) 067 [[arXiv:1105.0376](#)] [[INSPIRE](#)].
- [9] W. Altmannshofer, P. Paradisi and D.M. Straub, *Model-independent constraints on new physics in $b \rightarrow s$ transitions*, *JHEP* **04** (2012) 008 [[arXiv:1111.1257](#)] [[INSPIRE](#)].
- [10] C. Bobeth, G. Hiller, D. van Dyk and C. Wacker, *The decay $B \rightarrow K^*\ell^+\ell^-$ at low hadronic recoil and model-independent $\Delta B = 1$ constraints*, *JHEP* **01** (2012) 107 [[arXiv:1111.2558](#)] [[INSPIRE](#)].
- [11] F. Beaujean, C. Bobeth, D. van Dyk and C. Wacker, *Bayesian fit of exclusive $b \rightarrow s\bar{\ell}\ell$ decays: the standard model operator basis*, *JHEP* **08** (2012) 030 [[arXiv:1205.1838](#)] [[INSPIRE](#)].
- [12] W. Altmannshofer and D.M. Straub, *Cornering new physics in $b \rightarrow s$ transitions*, *JHEP* **08** (2012) 121 [[arXiv:1206.0273](#)] [[INSPIRE](#)].
- [13] N. Kosnik, *Model independent constraints on leptoquarks from $b \rightarrow s\ell^+\ell^-$ processes*, *Phys. Rev. D* **86** (2012) 055004 [[arXiv:1206.2970](#)] [[INSPIRE](#)].
- [14] T. Hurth and F. Mahmoudi, *The minimal flavour violation benchmark in view of the latest LHCb data*, *Nucl. Phys. B* **865** (2012) 461 [[arXiv:1207.0688](#)] [[INSPIRE](#)].
- [15] CDF collaboration, T. Aaltonen et al., *Search for $B_s^0 \rightarrow \mu^+\mu^-$ and $B^0 \rightarrow \mu^+\mu^-$ decays with CDF II*, *Phys. Rev. Lett.* **107** (2011) 191801 [Addendum *ibid.* **107** (2011) 239903] [[arXiv:1107.2304](#)] [[INSPIRE](#)].
- [16] LHCb collaboration, R. Aaij et al., *Strong constraints on the rare decays $B_s^0 \rightarrow \mu^+\mu^-$ and $B^0 \rightarrow \mu^+\mu^-$* , *Phys. Rev. Lett.* **108** (2012) 231801 [[arXiv:1203.4493](#)] [[INSPIRE](#)].
- [17] CMS collaboration, *Search for $B_s^0 \rightarrow \mu^+\mu^-$ and $B^0 \rightarrow \mu^+\mu^-$ decays*, *JHEP* **04** (2012) 033 [[arXiv:1203.3976](#)] [[INSPIRE](#)].
- [18] ATLAS collaboration, *Search for the decay $B_s^0 \rightarrow \mu^+\mu^-$ with the ATLAS detector*, *Phys. Lett. B* **713** (2012) 387 [[arXiv:1204.0735](#)] [[INSPIRE](#)].
- [19] J. Charles et al., *Predictions of selected flavour observables within the standard model*, *Phys. Rev. D* **84** (2011) 033005 [[arXiv:1106.4041](#)] [[INSPIRE](#)].

- [20] A.J. Buras, *Flavour visions*, PoS(BEAUTY 2011)008 [arXiv:1106.0998] [INSPIRE].
- [21] S. Descotes-Genon, J. Matias and J. Virto, *An analysis of $B_{d,s}$ mixing angles in presence of new physics and an update of $B_s \rightarrow \bar{K}^{0*}K^{0*}$* , *Phys. Rev. D* **85** (2012) 034010 [arXiv:1111.4882] [INSPIRE].
- [22] K. De Bruyn et al., *Branching ratio measurements of B_s decays*, *Phys. Rev. D* **86** (2012) 014027 [arXiv:1204.1735] [INSPIRE].
- [23] K. De Bruyn et al., *Probing new physics via the $B_s^0 \rightarrow \mu^+\mu^-$ effective lifetime*, *Phys. Rev. Lett.* **109** (2012) 041801 [arXiv:1204.1737] [INSPIRE].
- [24] F. Mahmoudi, S. Neshatpour and J. Orloff, *Supersymmetric constraints from $B_s \rightarrow \mu^+\mu^-$ and $B \rightarrow K^*\mu^+\mu^-$ observables*, *JHEP* **08** (2012) 092 [arXiv:1205.1845] [INSPIRE].
- [25] D. Becirevic, N. Kosnik, F. Mescia and E. Schneider, *Complementarity of the constraints on new physics from $B_s \rightarrow \mu^+\mu^-$ and from $B \rightarrow K\ell^+\ell^-$ decays*, *Phys. Rev. D* **86** (2012) 034034 [arXiv:1205.5811] [INSPIRE].
- [26] W. Altmannshofer et al., *Symmetries and asymmetries of $B \rightarrow K^*\mu^+\mu^-$ decays in the standard model and beyond*, *JHEP* **01** (2009) 019 [arXiv:0811.1214] [INSPIRE].
- [27] A. Khodjamirian, T. Mannel, A. Pivovarov and Y.-M. Wang, *Charm-loop effect in $B \rightarrow K^{(*)}\ell^+\ell^-$ and $B \rightarrow K^*\gamma$* , *JHEP* **09** (2010) 089 [arXiv:1006.4945] [INSPIRE].
- [28] P. Ball and R. Zwicky, *$B_{d,s} \rightarrow \rho, \omega, K^*, \phi$ decay form-factors from light-cone sum rules revisited*, *Phys. Rev. D* **71** (2005) 014029 [hep-ph/0412079] [INSPIRE].
- [29] A. Bharucha, T. Feldmann and M. Wick, *Theoretical and phenomenological constraints on form factors for radiative and semi-leptonic B-meson decays*, *JHEP* **09** (2010) 090 [arXiv:1004.3249] [INSPIRE].
- [30] F. Krüger and J. Matias, *Probing new physics via the transverse amplitudes of $B^0 \rightarrow K^{*0}(\rightarrow K^-\pi^+)\ell^+\ell^-$ at large recoil*, *Phys. Rev. D* **71** (2005) 094009 [hep-ph/0502060] [INSPIRE].
- [31] U. Egede, T. Hurth, J. Matias, M. Ramon and W. Reece, *New observables in the decay mode $\bar{B}_d \rightarrow \bar{K}^{*0}\ell^+\ell^-$* , *JHEP* **11** (2008) 032 [arXiv:0807.2589] [INSPIRE].
- [32] J. Matias, F. Mescia, M. Ramon and J. Virto, *Complete anatomy of $\bar{B}_d \rightarrow \bar{K}^{*0}(\rightarrow K\pi)\ell^+\ell^-$ and its angular distribution*, *JHEP* **04** (2012) 104 [arXiv:1202.4266] [INSPIRE].
- [33] U. Egede, T. Hurth, J. Matias, M. Ramon and W. Reece, *On the new physics reach of the decay mode $\bar{B}_d \rightarrow \bar{K}^{*0}\ell^+\ell^-$* , *JHEP* **10** (2010) 056 [arXiv:1005.0571] [INSPIRE].
- [34] D. Becirevic and E. Schneider, *On transverse asymmetries in $B \rightarrow K^*\ell^+\ell^-$* , *Nucl. Phys. B* **854** (2012) 321 [arXiv:1106.3283] [INSPIRE].
- [35] C. Bobeth, G. Hiller and D. van Dyk, *The benefits of $\bar{B} \rightarrow \bar{K}^*\ell^+\ell^-$ decays at low recoil*, *JHEP* **07** (2010) 098 [arXiv:1006.5013] [INSPIRE].
- [36] C. Bobeth, G. Hiller and G. Piranishvili, *CP asymmetries in $\bar{B} \rightarrow \bar{K}^*(\rightarrow \bar{K}\pi)\bar{\ell}\ell$ and untagged $\bar{B}_s, B_s \rightarrow \phi(\rightarrow K^+K^-)\bar{\ell}\ell$ decays at NLO*, *JHEP* **07** (2008) 106 [arXiv:0805.2525] [INSPIRE].
- [37] J. Matias, *On the S-wave pollution of $B \rightarrow K^*\ell^+\ell^-$ observables*, *Phys. Rev. D* **86** (2012) 094024 [arXiv:1209.1525] [INSPIRE].
- [38] M. Beneke, T. Feldmann and D. Seidel, *Systematic approach to exclusive $B \rightarrow V\ell^+\ell^-$, $V\gamma$ decays*, *Nucl. Phys. B* **612** (2001) 25 [hep-ph/0106067] [INSPIRE].
- [39] M. Beneke, T. Feldmann and D. Seidel, *Exclusive radiative and electroweak $b \rightarrow d$ and $b \rightarrow s$ penguin decays at NLO*, *Eur. Phys. J. C* **41** (2005) 173 [hep-ph/0412400] [INSPIRE].

- [40] T. Huber, E. Lunghi, M. Misiak and D. Wyler, *Electromagnetic logarithms in $\bar{B} \rightarrow X_s \ell^+ \ell^-$* , *Nucl. Phys. B* **740** (2006) 105 [[hep-ph/0512066](#)] [[INSPIRE](#)].
- [41] P. Gambino, M. Gorbahn and U. Haisch, *Anomalous dimension matrix for radiative and rare semileptonic B decays up to three loops*, *Nucl. Phys. B* **673** (2003) 238 [[hep-ph/0306079](#)] [[INSPIRE](#)].
- [42] M. Gorbahn and U. Haisch, *Effective Hamiltonian for non-leptonic $|\Delta F| = 1$ decays at NNLO in QCD*, *Nucl. Phys. B* **713** (2005) 291 [[hep-ph/0411071](#)] [[INSPIRE](#)].
- [43] C. Bobeth, P. Gambino, M. Gorbahn and U. Haisch, *Complete NNLO QCD analysis of $\bar{B} \rightarrow X_s \ell^+ \ell^-$ and higher order electroweak effects*, *JHEP* **04** (2004) 071 [[hep-ph/0312090](#)] [[INSPIRE](#)].
- [44] M. Misiak and M. Steinhauser, *NNLO QCD corrections to the $\bar{B} \rightarrow X_s \gamma$ matrix elements using interpolation in m_c* , *Nucl. Phys. B* **764** (2007) 62 [[hep-ph/0609241](#)] [[INSPIRE](#)].
- [45] M. Misiak et al., *Estimate of $\mathcal{B}(\bar{B} \rightarrow X_s \gamma)$ at $\mathcal{O}(\alpha_s^2)$* , *Phys. Rev. Lett.* **98** (2007) 022002 [[hep-ph/0609232](#)] [[INSPIRE](#)].
- [46] PARTICLE DATA GROUP collaboration, J. Beringer et al., *Review of particle physics*, *Phys. Rev. D* **86** (2012) 010001 [[INSPIRE](#)].
- [47] ALEPH, CDF, D0, DELPHI, L3, OPAL and SLD collaborations, J. Alcaraz, *Precision electroweak measurements and constraints on the standard model*, [arXiv:0911.2604](#) [[INSPIRE](#)].
- [48] C.W. Bauer, Z. Ligeti, M. Luke, A.V. Manohar and M. Trott, *Global analysis of inclusive B decays*, *Phys. Rev. D* **70** (2004) 094017 [[hep-ph/0408002](#)] [[INSPIRE](#)].
- [49] CKMFITTER GROUP collaboration, J. Charles et al., *CP violation and the CKM matrix: assessing the impact of the asymmetric B factories*, *Eur. Phys. J. C* **41** (2005) 1 [[hep-ph/0406184](#)] [[INSPIRE](#)] and online update on <http://ckmfitter.in2p3.fr/>.
- [50] A.L. Kagan and M. Neubert, *Isospin breaking in $B \rightarrow K^* \gamma$ decays*, *Phys. Lett. B* **539** (2002) 227 [[hep-ph/0110078](#)] [[INSPIRE](#)].
- [51] C. Davies, *Standard model heavy flavor physics on the lattice*, [PoS\(Lattice 2011\)019](#) [[arXiv:1203.3862](#)] [[INSPIRE](#)].
- [52] H.H. Asatrian, H.M. Asatrian, C. Greub and M. Walker, *Two loop virtual corrections to $B \rightarrow X_s \ell^+ \ell^-$ in the standard model*, *Phys. Lett. B* **507** (2001) 162 [[hep-ph/0103087](#)] [[INSPIRE](#)].
- [53] S. Descotes-Genon, J. Matias and J. Virto, *New physics constraints from optimized observables in $B \rightarrow K^* \ell \ell$ at large recoil*, *AIP Conf. Proc.* **1492** (2012) 103 [[arXiv:1209.0262](#)] [[INSPIRE](#)].
- [54] D. Becirevic, E. Kou, A. Le Yaouanc and A. Tayduganov, *Future prospects for the determination of the Wilson coefficient $C'_{7\gamma}$* , *JHEP* **08** (2012) 090 [[arXiv:1206.1502](#)] [[INSPIRE](#)].
- [55] HEAVY FLAVOR AVERAGING GROUP collaboration, D. Asner et al., *Averages of b -hadron, c -hadron and τ -lepton properties*, [arXiv:1010.1589](#) [[INSPIRE](#)].
- [56] T. Huber, T. Hurth and E. Lunghi, *Logarithmically enhanced corrections to the decay rate and forward backward asymmetry in $\bar{B} \rightarrow X_s \ell^+ \ell^-$* , *Nucl. Phys. B* **802** (2008) 40 [[arXiv:0712.3009](#)] [[INSPIRE](#)].
- [57] S. Descotes-Genon, J. Matias, M. Ramon and J. Virto, *Implications from clean observables for the binned analysis of $B \rightarrow K^* \mu^+ \mu^-$ at large recoil*, [arXiv:1207.2753](#) [[INSPIRE](#)].