

INTERSTELLAR EXTINCTION AND POLARIZATION IN THE TAURUS DARK CLOUDS: THE OPTICAL PROPERTIES OF DUST NEAR THE DIFFUSE/DENSE CLOUD INTERFACE

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ABSTRACT

Observations of interstellar linear polarization in the spectral range 0.35–2.2 μm are presented for several stars reddened by dust in the Taurus region. Combined with a previously published study by Whittet et al., these results represent the most comprehensive data set available on the spectral dependence of interstellar polarization in this nearby dark cloud (a total of 27 sight lines). Extinction data for these and other reddened stars in Taurus are assembled for the same spectral range, combining published photometry and spectral classifications with photometry from the Two Micron All Sky Survey. The polarization and extinction curves are characterized in terms of the parameters λ_{max} (the wavelength of maximum polarization) and R_V (the ratio of total to selective extinction), respectively. The data are used to investigate in detail the question of whether the optical properties of the dust change systematically as a function of environment, considering stars observed through progressively more opaque (and thus progressively denser) regions of the cloud. At low visual extinctions ($0 < A_V < 3$), the dust has $R_V = 2.97 \pm 0.15$, implying optical properties closely similar to those of “normal” dust in the diffuse interstellar medium (ISM). However, λ_{max} is significantly higher than the diffuse-ISM average toward some stars in this extinction range, a result which we attribute to size-dependent failure of the grain alignment mechanism. For extinctions $A_V > 3$, real changes in grain properties occur, characterized by observed R_V values in the range 3.5–4.0. A simple model for the development of R_V with A_V suggests that R_V may approach values of 4.5 or more in the densest regions of the cloud. The transition between “normal” extinction and “dense cloud” extinction occurs at $A_V \sim 3.2$, a value coincident with the threshold extinction above which H_2O -ice is detected on grains within the cloud. Changes in R_V are thus either a direct consequence of mantle growth or occur under closely similar physical conditions. Dust in Taurus appears to be in a different evolutionary state compared with other nearby dark clouds, such as ρ Oph, in which coagulation is the dominant physical process.

Subject headings: dust, extinction — ISM: individual (Taurus Dark Cloud) — ISM: magnetic fields — polarization

1. INTRODUCTION

The extensive complex of dark clouds in the Taurus region (hereafter TDC) provides an excellent opportunity to study the properties of interstellar matter under physical conditions close to the interface between diffuse atomic and dense molecular phases. At a distance of 140 pc (Elias 1978a; Kenyon, Dobrzyzcka, & Hartmann 1994), the TDC is one of the closest regions of current star formation to the solar system. Little foreground or background dust appears to be present, virtually all of the extinction and reddening of objects within and behind the cloud occurring in the cloud itself (Straižys & Meištas 1980; Meištas & Straižys 1981). Studies of the 3 μm H_2O -ice absorption feature and its correlation with visual extinction have shown that icy grain material is present only at extinctions above some threshold; i.e., the dust grains acquire detectable ice mantles at $A_V > A_{\text{th}}$, where $A_{\text{th}} \approx 3$ is the threshold extinction in magnitudes (Whittet et al. 1983; 1988; Smith, Sellgren & Brooke 1993; Teixeira & Emerson 1999; Murakawa, Tamura, &

Nagata 2000). The value of A_{th} varies from cloud to cloud, and is generally lower in the TDC compared with other dark clouds studied to date (e.g., Whittet et al. 1996 and references therein), consistent with an absence of internal sources of desorbing radiation (Williams, Hartquist, & Whittet 1992). In this paper, we take the threshold extinction to be a marker of the transition between the diffuse outer layers of the cloud and the cold, dense interior. It is likely that the ice threshold is quite closely related to the phase change between predominantly atomic and predominantly molecular gas. Ice mantles are predicted to form primarily by surface reactions between accreted H and O atoms (Jones & Williams 1984), and thus grow most efficiently on grains immersed in the densest gas that has a significant atomic H content, i.e., close to the $\text{H I}/\text{H}_2$ interface. Note also that millimeter-wave emission from gas phase CO, the primary observational tracer of H_2 , is likewise detected only above a threshold extinction that is quantitatively similar to the ice threshold in the TDC (Frerking, Langer, & Wilson 1982).

Several physical processes can lead to changes in the optical (extinction and polarization) properties of dust in dense regions of dark clouds. Mantle growth, in which atoms and molecules accrete from the gas to form icy coatings, has already been mentioned. The rate of growth is expected to be independent of the size of the particle in the absence of desorbing radiation, and the appearance of

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mantles should thus introduce a systematic shift of the entire size distribution toward larger sizes. Grains may also grow by grain-grain collisions that lead to coagulation, and this will tend to redistribute the size distribution as small grains become depleted by inclusion in larger ones. It is the latter process that is thought to be primarily responsible for systematic changes observed in the extinction curve (e.g., Cardelli, Clayton, & Mathis 1989; Mathis 1996). Finally, when polarization is considered, variations in the efficiency of alignment by the ambient magnetic field must also be taken into account. If these changes are size-dependent, they will affect the spectral dependence as well as the degree of polarization.

Whereas the *distribution* of reddening and polarizing material in the TDC has been studied quite extensively (e.g., Straizys & Meistas 1980; Meistas & Straizys 1981; Moneti et al. 1984; Arce & Goodman 1999), the *wavelength-dependence* of these phenomena has been investigated only for a relatively few lines of sight. Available data for stars with low to moderate extinctions (Vrba & Rydgren 1985; Whittet et al. 1992; Kenyon et al. 1994) hint that the dust has generally “normal” optical properties, characterized by values of the parameters $R_V = A_V/E_{B-V}$ (the ratio of total to selective extinction) and λ_{\max} (the wavelength of maximum linear polarization) close to mean values for the general interstellar medium (ISM), i.e., $R_V \approx 3.1$ and $\lambda_{\max} \approx 0.55 \mu\text{m}$. However, objects with deep ice absorption features are generally not accessible to observation in the visible because of their high extinctions. The best studied case of a visible star with detectable ice absorption is HD 29647: its extinction curve has a somewhat elevated value of R_V (~ 3.6 ; Crutcher 1985; Vrba & Rydgren 1985; Straizys, Cernis & Hayes 1985), and displays more spectacular departures from average behavior in the ultraviolet (Cardelli & Savage 1988).

The main aim of this paper is a detailed study of the wavelength dependences of interstellar polarization and extinction in the TDC, based on a greatly extended sample of sight lines compared with previous work. In particular, we focus on the question of whether the optical properties of the dust change systematically as a function of environment near the ice threshold. Our data set, described in § 2, includes visible and infrared broadband polarimetry, obtained with the United Kingdom Infrared Telescope, and infrared photometry extracted from the Two Micron All Sky Survey. The data are analyzed and discussed in § 3. A simple model for the dependence of R_V on visual extinction is described and compared with observations in § 4. Constraints on the relative importance of mantle growth and coagulation as physical processes leading to changes in R_V are discussed in § 5. A summary of the principal results and suggestions for further research appear in the final section.

2. THE DATA SET

A total of 54 reddened stars in the general direction of the TDC complex are included in this study. The primary objectives of the observing program were (1) to extend the existing polarimetric data set for Taurus stars reported by Whittet et al. (1992) and (2) to determine reliable extinction parameters for stars with polarimetric data available. The final sample contains 27 targets with both polarization and extinction information. Most of our targets were selected from lists of visually bright ($V < 12$ mag) reddened and/or polarized stars (e.g., Straizys & Meistas 1980; Moneti et al.

1984).⁴ In addition to visually-selected targets, we include three objects from the Elias (1978a) near infrared survey of the TDC. To increase the sample size for extinction studies, we also present a compilation of extinction data for a further 27 stars that lack spectropolarimetric data, selected from Vrba & Rydgren (1985), Kenyon et al. (1994), and Arce & Goodman (1999).⁵ All objects included, both in the polarimetry observing program and in the additional set of reddened stars, are thought to be *field stars*, i.e., situated behind the TDC complex and reddened by foreground dust within it. Embedded young stellar objects are specifically excluded, owing to uncertainties in their intrinsic colors and the possibility of circumstellar polarization or infrared emission.

2.1. Polarimetry

New observations of interstellar linear polarization were made for a total of 11 program stars during an observing run with the United Kingdom Infrared Telescope (UKIRT) at Mauna Kea Observatory, Hawaii, on the nights of 1995 January 6–8. The instrument used was the Hatfield polarimeter (Hatpol), described in detail by Hough, Peacock, & Bailey (1991). Simultaneous observations were possible in six broadband filters (U, B, V, R, I , and one selected from J, H, K). The effective wavelengths of the filters were 0.36, 0.43, 0.55, 0.63, 0.78, 1.25, 1.65 and $2.20 \mu\text{m}$, respectively. The data collection and reduction procedures were closely similar to those described in our previously published studies of interstellar polarization using this instrument (Whittet et al. 1992, 1994). Resulting values of the degree (p_λ) and position angle (θ_λ) of polarization, and their associated uncertainties, are presented in Table 1.

It is well known that the wavelength dependence of linear polarization can be described to a close approximation by the empirical formula known as the Serkowski law:

$$p_\lambda/p_{\max} = \exp[-K \ln^2(\lambda_{\max}/\lambda)] \quad (1)$$

where p_{\max} is the peak percentage polarization and λ_{\max} the wavelength corresponding to p_{\max} (Serkowski 1973; Serkowski, Mathewson & Ford 1975). With K treated as a free parameter, following Wilking et al. (1980) and Wilking, Lebofsky, & Rieke (1982), least-squares fits of equation (1) to data for stars with a range of λ_{\max} show that K and λ_{\max} are linearly correlated:

$$K = c_1 \lambda_{\max} + c_2 \quad (2)$$

For a variety of interstellar environments and physical conditions, the average values of the slope and intercept are $c_1 = 1.66 \pm 0.09$ and $c_2 = 0.01 \pm 0.05$ (Whittet et al. 1992).

⁴ Visually-selected objects are designated by their Henry Draper (HD) numbers throughout this paper. HD numbers for Taurus stars in the main Henry Draper Catalog have five digits, those in the Henry Draper Extension have six digits. Some authors, notably Kenyon et al. (1994), use SAO numbers in preference to HD numbers: the correspondence may be easily found, e.g., with reference to the Simbad database.

⁵ Arce & Goodman (1999) present reddening data for some 100 stars in the direction of the TDC. However, the vast majority have relatively low extinction ($A_V < 3$), and their spectral classifications do not include luminosity classes; hence the uncertainties in the color excesses and in A_V and R_V are larger than for typical stars in our sample. We selected for analysis the only stars (091085 and 141134) in the Arce & Goodman compilation that have extinctions substantially above the ice threshold. Because of their relatively high extinctions, errors introduced by uncertainties in spectral types are less important.

TABLE 1
RESULTS OF BROADBAND POLARIMETRY

Star	Filter	p_λ (%)	θ_λ (deg)	
HD 28225	U	1.53 ± 0.05	13.6 ± 0.8	
	B	1.68 ± 0.03	12.5 ± 0.5	
	V	1.93 ± 0.03	12.4 ± 0.4	
	R	1.78 ± 0.06	13.0 ± 0.9	
	I	1.58 ± 0.06	12.0 ± 1.0	
	J	0.98 ± 0.04	12.9 ± 1.4	
	H	0.54 ± 0.03	15.8 ± 1.7	
	K	0.34 ± 0.02	10.9 ± 1.5	
HD 283367	U	0.81 ± 0.03	133.4 ± 1.2	
	B	1.05 ± 0.04	127.8 ± 1.2	
	V	1.35 ± 0.03	126.8 ± 0.7	
	R	1.41 ± 0.04	127.7 ± 0.8	
	I	1.36 ± 0.04	124.3 ± 0.8	
	J	0.97 ± 0.05	119.5 ± 1.6	
	H	0.57 ± 0.03	122.6 ± 2.3	
	K	0.36 ± 0.04	115.8 ± 3.5	
	HD 283642	U	1.59 ± 0.08	36.5 ± 1.3
		B	1.71 ± 0.02	34.2 ± 0.4
V		1.98 ± 0.02	32.1 ± 0.3	
R		1.97 ± 0.02	32.0 ± 0.3	
I		1.82 ± 0.03	29.0 ± 0.5	
J		1.14 ± 0.02	28.0 ± 0.7	
H		0.74 ± 0.05	26.1 ± 1.7	
K		0.42 ± 0.03	28.4 ± 2.5	
HD 283643		U	1.12 ± 0.06	33.1 ± 1.4
		B	1.16 ± 0.03	29.5 ± 0.7
	V	1.31 ± 0.03	24.7 ± 0.7	
	R	1.30 ± 0.03	25.5 ± 0.6	
	I	1.40 ± 0.04	21.8 ± 0.8	
	J	0.81 ± 0.03	18.8 ± 1.2	
	H	0.55 ± 0.04	18.9 ± 2.1	
	K	0.28 ± 0.06	(12 ± 7)	
	HD 283757	U	2.35 ± 0.05	44.2 ± 0.7
		B	2.46 ± 0.03	41.7 ± 0.3
V		2.95 ± 0.02	40.2 ± 0.2	
R		2.77 ± 0.02	40.0 ± 0.2	
I		2.73 ± 0.03	40.2 ± 0.3	
J		1.56 ± 0.04	37.0 ± 1.0	
H		0.97 ± 0.05	39.8 ± 1.4	
K		0.67 ± 0.09	38.7 ± 3.5	
HD 283815		U	2.40 ± 0.06	45.3 ± 0.7
		B	2.47 ± 0.02	45.9 ± 0.3
	V	2.82 ± 0.03	45.2 ± 0.3	
	R	2.78 ± 0.03	44.9 ± 0.3	
	I	2.58 ± 0.03	45.7 ± 0.4	
	J	1.49 ± 0.03	46.1 ± 0.8	
	H	0.92 ± 0.04	49.0 ± 1.8	
	K	0.57 ± 0.04	44.2 ± 2.5	
	HD 283877	U	1.23 ± 0.03	29.6 ± 0.7
		B	1.33 ± 0.03	28.1 ± 0.6
V		1.68 ± 0.04	27.2 ± 0.6	
R		1.59 ± 0.03	25.7 ± 0.5	
I		1.53 ± 0.02	28.2 ± 0.5	
J		1.00 ± 0.04	25.3 ± 1.4	
H		0.67 ± 0.06	(7 ± 5)	
K		0.36 ± 0.12	(14 ± 7)	
HD 283879		U	3.12 ± 0.08	33.9 ± 0.8
		B	3.43 ± 0.04	31.5 ± 0.3
	V	4.16 ± 0.03	29.6 ± 0.2	
	R	4.07 ± 0.03	29.0 ± 0.2	
	I	3.95 ± 0.04	29.1 ± 0.2	
	J	2.36 ± 0.04	27.3 ± 0.5	
	H	1.39 ± 0.03	27.8 ± 0.7	
	K	0.86 ± 0.05	28.4 ± 1.7	

TABLE 1—Continued

Star	Filter	p_λ (%)	θ_λ (deg)
Elias 3	V	2.37 ± 0.40	107.6 ± 4.7
	R	2.45 ± 0.13	107.1 ± 1.5
	I	2.35 ± 0.08	108.6 ± 0.9
	J	1.72 ± 0.04	108.7 ± 0.7
	H	1.09 ± 0.04	109.6 ± 1.0
	K	0.65 ± 0.02	108.4 ± 0.7
	Elias 29	U	3.2 ± 1.3
B		4.22 ± 0.16	31.0 ± 1.1
V		4.37 ± 0.07	32.2 ± 0.4
R		4.39 ± 0.04	31.3 ± 0.3
I		3.91 ± 0.06	30.4 ± 0.3
Elias 30	J	2.37 ± 0.02	29.8 ± 0.3
	H	1.49 ± 0.03	29.6 ± 0.7
	B	4.46 ± 0.36	36.8 ± 2.1
	V	5.57 ± 0.07	35.3 ± 0.3
	R	5.31 ± 0.05	34.0 ± 0.2
	I	4.94 ± 0.04	33.9 ± 0.2
	J	2.86 ± 0.06	32.4 ± 0.7
	H	1.81 ± 0.05	33.1 ± 0.9
	K	1.14 ± 0.04	32.1 ± 1.3

The dependence of K on λ_{\max} , known as the Wilking law, implies a systematic decrease in the width of the polarization curve (proportional to $1/K$) with increasing λ_{\max} . The simplest explanation of this behavior is a systematic reduction in the relative numbers of small, aligned grains in regions of high λ_{\max} (Clayton et al. 1995; Martin, Clayton, & Wolff 1999).

Fits to the observed values of p_λ were calculated, using equation (1) with K treated as a free parameter. The values of p_{\max} , λ_{\max} , and K generated by the fits are listed in Table 2. The fitting process was applied to all stars that satisfy our selection criteria and with Hatpol data available, either from Table 1 of this paper or from our previously published results (Whittet et al. 1992). The fit parameters given in Table 2 are generally consistent with those listed in Whittet et al. (1992) for stars in common (Elias 29 is an exception: the previous observations of this star are of much lower quality than those presented here). Also given in Table 2 are the mean, $\langle\theta\rangle$, and standard deviation, $\sigma(\theta)$, of the position angle for each star. Values of θ_λ with large uncertainty (in parentheses in Table 1) were excluded from the calculation. A few stars show evidence for position-angle rotation with wavelength, i.e., they exhibit a systematic trend in θ_λ with passband, and thus have values of $\sigma(\theta)$ appreciably greater than the mean observational error. We take $\sigma(\theta) > 3^\circ$ as indicative of significant rotation: five stars (HD 29647, 283367, 283642, 283643, and 283809) fall into this group.

As a consistency check, we plot the correlation of λ_{\max} with K (Table 2) in Figure 1. A linear least-squares fit to the 27 data points (shown as the solid line in Fig. 1) has slope and intercept $c_1 = 1.5 \pm 0.4$ and $c_2 = 0.1 \pm 0.1$ (eq. [2]). We conclude that the correlation for Taurus is not significantly different from the average relationship adopted by Whittet et al. (1992), shown as the dashed line in Figure 1.

2.2. Infrared Photometry

Previous studies of interstellar extinction in the TDC were hindered by a sparsity of infrared data needed to determine R_V and A_V (§ 2.3). This problem is now alleviated by the availability of results from the Two Micron All Sky

TABLE 2
SERKOWSKI FIT PARAMETERS AND MEAN POSITION ANGLES

Star	p_{\max} (%)	λ_{\max} (μm)	K	$\langle\theta\rangle$ (deg)	$\sigma(\theta)$ (deg)	References
HD 28170	1.92 ± 0.03	0.55 ± 0.01	0.98 ± 0.06	89.4	0.7	1
HD 28225	1.88 ± 0.04	0.58 ± 0.02	1.05 ± 0.09	12.9	1.3	2
HD 28975	3.17 ± 0.03	0.54 ± 0.01	0.89 ± 0.05	57.6	2.3	1
HD 29333	5.25 ± 0.07	0.54 ± 0.02	0.89 ± 0.07	70.9	1.5	1
HD 29647	2.30 ± 0.02	0.73 ± 0.01	1.15 ± 0.04	71.4	7.3	1
HD 29835	4.07 ± 0.07	0.50 ± 0.03	0.93 ± 0.11	23.0	1.5	1
HD 30168	4.07 ± 0.09	0.53 ± 0.03	0.83 ± 0.09	24.6	0.9	1
HD 30675	3.90 ± 0.03	0.52 ± 0.01	0.99 ± 0.05	58.7	1.3	1
HD 279652.....	1.29 ± 0.09	0.47 ± 0.09	0.78 ± 0.26	89.7	2.3	1
HD 279658.....	2.81 ± 0.03	0.53 ± 0.01	0.99 ± 0.06	146.0	1.4	1
HD 283367.....	1.43 ± 0.05	0.71 ± 0.02	1.23 ± 0.11	124.7	5.1	2
HD 283637.....	2.72 ± 0.07	0.57 ± 0.05	0.89 ± 0.19	6.5	0.8	1
HD 283642.....	2.01 ± 0.03	0.63 ± 0.01	1.12 ± 0.08	30.8	3.3	2
HD 283643.....	1.35 ± 0.03	0.64 ± 0.02	1.01 ± 0.12	24.6	4.9	2
HD 283701.....	3.18 ± 0.04	0.60 ± 0.02	0.89 ± 0.07	33.0	1.3	1
HD 283725.....	4.84 ± 0.07	0.50 ± 0.02	0.79 ± 0.06	66.4	1.2	1
HD 283757.....	2.92 ± 0.02	0.62 ± 0.01	1.05 ± 0.06	40.2	2.0	2
HD 283800.....	3.95 ± 0.11	0.53 ± 0.02	0.86 ± 0.14	27.0	0.9	1
HD 283809.....	6.70 ± 0.10	0.59 ± 0.02	0.97 ± 0.10	52.4	4.1	1
HD 283812.....	6.30 ± 0.09	0.55 ± 0.02	0.96 ± 0.07	31.8	1.3	1
HD 283815.....	2.86 ± 0.04	0.61 ± 0.01	1.11 ± 0.06	45.8	1.3	2
HD 283855.....	5.13 ± 0.09	0.51 ± 0.01	0.91 ± 0.07	46.0	0.6	1
HD 283877.....	1.65 ± 0.04	0.65 ± 0.01	1.01 ± 0.11	27.4	1.5	2
HD 283879.....	4.24 ± 0.04	0.65 ± 0.01	1.22 ± 0.04	29.6	2.0	2
Elias 3	2.48 ± 0.03	0.69 ± 0.02	1.03 ± 0.07	108.3	0.8	2
Elias 29	4.51 ± 0.08	0.59 ± 0.02	1.12 ± 0.10	30.7	0.9	2
Elias 30	5.49 ± 0.05	0.64 ± 0.01	1.24 ± 0.05	33.9	1.5	2

REFERENCES.—(1) Whittet et al. 1992. (2) This paper, Table 1.

Survey (2MASS): the first and second releases (1999 May and 2000 March) together cover almost all of the Taurus region. 2MASS provides a database of stellar photometry in passbands centered at $1.25 \mu\text{m}$ (J), $1.65 \mu\text{m}$ (H) and $2.17 \mu\text{m}$ (K_s).⁶ In general, any differences between 2MASS filter characteristics and those of other near infrared photometric systems in common use (e.g., Bessell & Brett 1988) should

⁶ See the 2MASS World Wide Web site at <http://pegasus.phast.umass.edu> for full details.

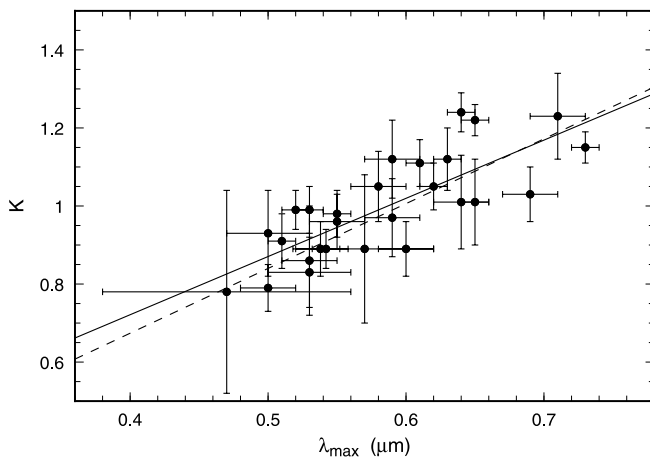


FIG. 1.—Plot of the Serkowski “width” parameter K against λ_{\max} , using data from Table 2. The solid line is a linear least-squares fit to the points. The dashed line is the linear relationship adopted by Whittet et al. (1992) as representative of the interstellar medium over a range of environments.

be very small. The internal calibration of 2MASS photometry is thought to be extremely precise. Relationships between 2MASS and other photometric systems are currently under investigation and not yet firmly established. An important concern in this paper is to match observed and intrinsic colors.

We compared 2MASS photometry with previous JHK photometry, as available, for stars in the Taurus region: the data are listed in Table 3. 2MASS magnitudes are presented as they appear in the database without rounding. Nominal errors in individual 2MASS J , H , and K magnitudes are typically ± 0.03 , and in no case do they exceed ± 0.05 . In addition to results from previous literature, the comparison photometry in Table 3 includes previously unpublished UKIRT observations for 12 stars. The latter were obtained on the night of 1988 December 15, using the solid nitrogen-cooled infrared spectrophotometer UKT6 and standard JHK passbands. Observational uncertainties in these magnitudes are typically ± 0.02 or less. Standard stars for zero-point calibration were selected from Elias et al. (1982), and results are on the CIT photometric system.

Figures 2 and 3 show correlation plots (2MASS vs. literature) for the two color indices most relevant to this study, i.e., $V-K$ and $J-K$. The V magnitudes used and their references are listed in Tables 4 and 5. Infrared data of different authors are distinguished by plotting symbol; all literature values have been transformed to the CIT system as necessary.⁷ Also shown in Figures 2 and 3 are linear least-squares fits to the complete data set in each case. The

⁷ Infrared photometry from Vrba & Rydgren (1985) and Straižys et al. (1985) is on the Kitt Peak system, which is similar to the Johnson system.

TABLE 3
COMPARISON OF 2MASS AND LITERATURE *JHK* PHOTOMETRY

STAR	2MASS			LITERATURE ^a			REFERENCES ^b
	<i>J</i>	<i>H</i>	<i>K</i>	<i>J</i>	<i>H</i>	<i>K</i>	
HD 27405.....	7.290	7.294	7.275	7.32	...	7.31	1
HD 27778A.....	5.900	5.892	5.902	5.96	...	5.94	1
HD 28149.....	5.642	5.717	5.731	5.70	...	5.76	1
HD 28170.....	7.695	7.576	7.463	7.76	7.57	7.49	2
HD 28975.....	7.415	7.210	7.097	7.47	7.24	7.13	3
HD 28991.....	7.058	6.994	6.845	7.08	...	6.89	1
HD 29333.....	7.140	6.959	6.824	7.16	6.96	6.85	2
HD 29647.....	5.932	5.578	5.359	5.98	...	5.41	1
				5.92	5.57	5.36	3
HD 29835.....	5.754	...	4.906	5.78	5.11	4.96	2
HD 30168.....	6.882	6.865	6.773	6.94	6.85	6.81	2
HD 30675.....	6.707	6.687	6.641	6.79	6.71	6.69	2
HD 279652.....	8.941	8.840	8.785	8.92	8.80	8.74	2
HD 279658.....	8.590	8.458	8.458	8.63	8.49	8.44	2
HD 283367.....	8.954	8.780	8.630	8.99	8.77	8.67	3
HD 283550.....	7.204	7.057	6.955	7.22	...	6.97	1
HD 283637.....	9.522	9.285	9.190	9.50	9.28	9.16	2
				9.57	9.32	9.18	3
HD 283701.....	7.792	7.561	7.435	7.81	7.60	7.48	2
HD 283725.....	7.912	7.599	7.425	7.93	7.57	7.46	2
HD 283740.....	9.547	9.454	9.405	9.44	...	9.32	1
HD 283800.....	8.845	8.751	8.698	8.83	8.73	8.68	2
HD 283809.....	6.967	6.455	6.162	7.00	6.44	6.14	4
HD 283855.....	9.517	9.349	9.243	9.47	9.28	9.20	2
VR c.....	7.841	7.585	7.364	7.94	7.60	7.42	3
VR d.....	10.188	9.960	9.837	10.25	10.00	9.85	3
VR e.....	9.606	9.351	9.187	9.61	9.33	9.17	3
Elias 4.....	7.577	6.858	6.673	7.57	6.90	6.73	5

^a CIT photometric system unless otherwise indicated.

^b References: 1, 2, 3, and 6 are as in Table 4; (12) Arce & Goodman 1999.

correlation coefficients are 0.999 and 0.994 for $V-K$ and $J-K$, respectively. The fits result in the following transformations:

$$(V-K)_{\text{CIT}} = (1.000 \pm 0.005)(V-K)_{2\text{MASS}} - (0.013 \pm 0.005) \quad (3)$$

and

$$(J-K)_{\text{CIT}} = (0.94 \pm 0.01)(J-K)_{2\text{MASS}} + (0.02 \pm 0.01) \quad (4)$$

Our transformation for $J-K$ (eq. [4]) is consistent with the results of an independent study by J. Carpenter (2000, private communication). These equations were used to homogenize the photometry (effectively tying 2MASS to CIT) prior to calculation of color excesses.

2.3. Color Excesses and Extinction Results

A catalog of spectral classifications, color excesses, visual extinctions and R_V values for our polarimetry program stars is presented in Table 4, together with bibliographic information. Corresponding data for additional reddened stars (lacking polarimetry) appears in Table 5. Intrinsic $B-V$ colors used to calculate E_{B-V} are taken from Schmidt-Kaler (1982). E_{V-K} and E_{J-K} were calculated using infrared photometry from 2MASS, averaged with data from other sources (Table 3) where available, all data being tied to the CIT photometric system (§ 2.2). Intrinsic $V-K$ and $J-K$ colors were taken from Bessell & Brett

(1988) and Glass (1997), transformed to the CIT system using equations given by these authors.⁸ Luminosity class V was assumed for stars lacking two-dimensional classifications.

Determination of R_V and A_V for individual stars is based on the assumption

$$R_V = \left[\frac{E_{V-\lambda}}{E_{B-V}} \right]_{\lambda \rightarrow \infty} \quad (5)$$

(e.g., Whittet 1992). This extrapolation is expected to hold unless there is a substantial population of very large grains that give neutral (wavelength-independent) extinction, a possibility that seems inconsistent with observational evidence and abundance constraints (Whittet 1992). In practice, λ in equation (5) is set to the longest-wavelength infrared passband for which photometry is available. An approximation consistent with theoretical models,

$$R_V \approx 1.1 \frac{E_{V-K}}{E_{B-V}} \quad (6)$$

(Whittet & van Breda 1978), is adopted here, in common with many other investigations of interstellar extinction. The error in R_V for a given star depends primarily on the uncertainty of its spectral type and on the degree of extinc-

⁸ Note that Glass (1997) includes early (O-B8) spectral types not covered by the Bessell & Brett (1988) compilation.

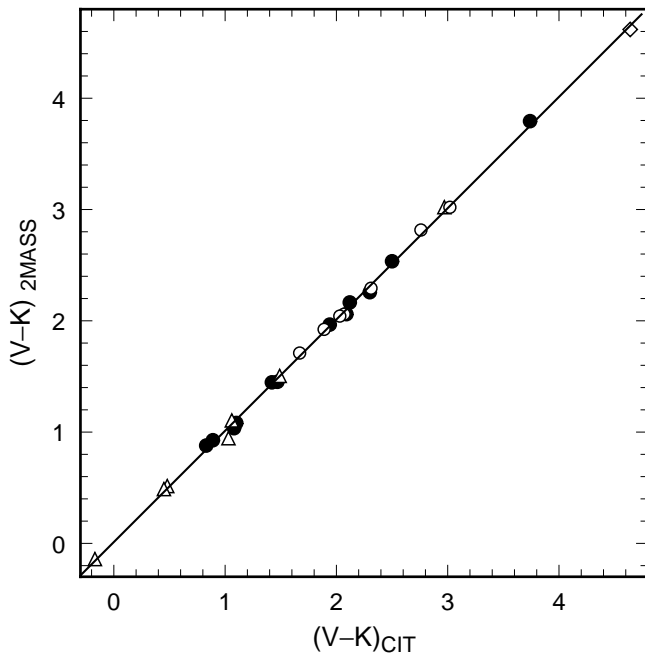


FIG. 2.—Correlation of $V-K$ colors based on 2MASS and literature infrared photometry for stars listed in Table 3. Literature values are from this paper (*solid circles*), Vrba & Rydgren (1985; *open circles*); Kenyon et al. (1994; *upright triangles*); Straizys et al. (1985; *diamonds*) and Elias (1978a; *inverted triangles*). All literature values have been transformed to the CIT system. The unweighted linear least-squares fit to all data points is shown.

tion; ± 0.15 is a representative value. It follows from equation (6) that

$$A_V \approx 1.1E_{V-K}. \quad (7)$$

This relation was used to calculate A_V for most individual stars in our study.

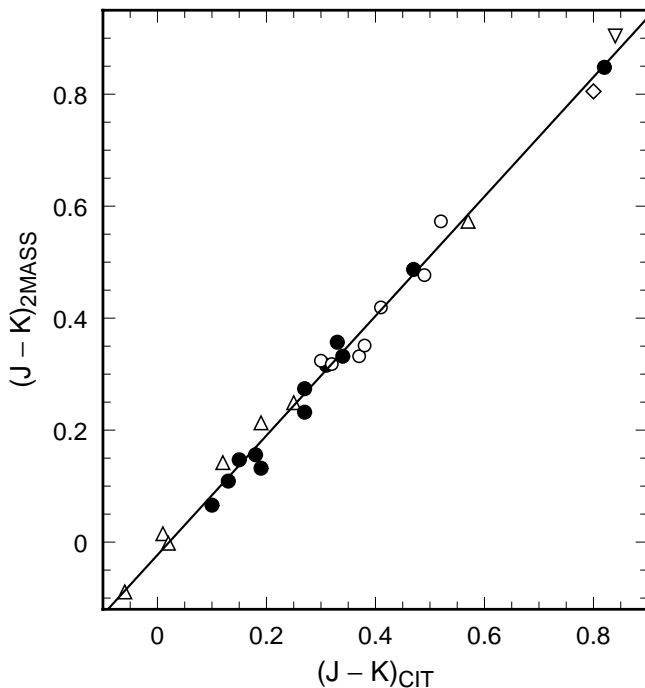


FIG. 3.—Same as Fig. 2, but for $J-K$

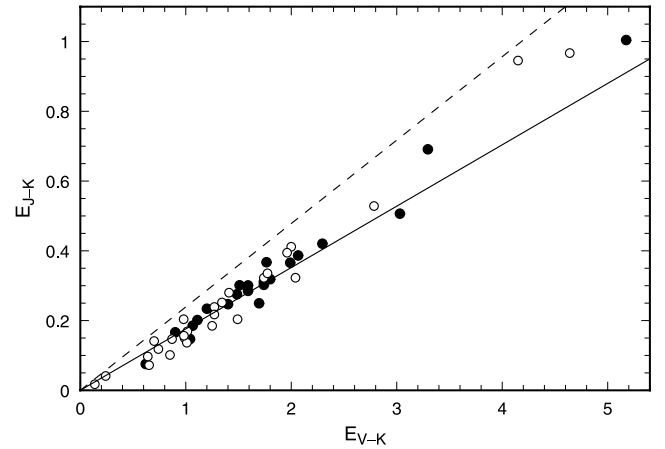


FIG. 4.—Plot of E_{J-K} vs. E_{V-K} . Filled circles denote stars from the polarimetry sample (Table 4) and open circles denote other reddened stars (Table 5). The solid line is the linear least-squares fit through the origin to all points with $E_{J-K} < 0.6$, with slope equivalent to $r = A_V/E_{J-K} = 6.2$. The dashed line is the extinction law ($r = 4.6$) adopted by Elias (1978a).

Estimates of A_V for reddened stars that lack visual photometry (of which there are four in Table 4) can be made by extrapolation from an infrared color excess such as E_{J-K} :

$$A_V = rE_{J-K}, \quad (8)$$

where r is the ratio of total to selective near-infrared extinction. The value of r depends on the form of the interstellar extinction curve over the relevant spectral range. Elias (1978a) adopted $r = 4.6$ to calculate A_V toward field stars detected in his $2 \mu\text{m}$ survey of the TDC—a value actually determined from the extinction curve of the star S-1 in the ρ Ophiuchi dark cloud (Elias 1978b). However, Smith, Sellgren, & Brooke (1993) propose $r \approx 5.4$ in the TDC, a value closer to that determined for diffuse (low-density) regions of the ISM ($r \approx 5.9$; Torres-Dodgen, Tapia, & Carrol 1991; He et al. 1995 and references therein).

The plot of E_{J-K} versus E_{V-K} for our sample (Fig. 4) provides constraints on the value of r appropriate to dust in the TDC. This correlation is predicted to be a straight line through the origin of slope $1.1/\langle r \rangle$ for dust with uniform optical properties. The best linear fit through the origin to all data points with $E_{J-K} < 0.6$ (*solid line*) yields $\langle r \rangle = 6.2 \pm 0.4$, a value indistinguishable from that for the diffuse-ISM. Note, however, that some of the most highly reddened stars in Figure 4 tend to have lower r , and thus plot closer to the Elias extinction law ($r = 4.6$; *dashed line*). We calculated A_V for stars lacking B and V photometry in Table 4 using equation (8) with $r = 6$ ($E_{J-K} < 0.6$) and $r = 5.3$ ($E_{J-K} > 0.6$). The latter value represents a compromise between the Elias and diffuse-ISM extinction laws; it will tend to overestimate the extinction toward stars with the highest color excesses (e.g., Elias 13) if r decreases systematically with increasing A_V in the inner regions of the cloud.

3. ANALYSIS

3.1. The Threshold Extinction

As discussed in § 1, the threshold extinction (A_{th}) for ice absorption is a key parameter in our characterization of the diffuse/dense cloud interface. A_{th} is determined from the correlation of A_V with the optical depth ($\tau_{3.0}$) of the $3 \mu\text{m}$ ice

TABLE 4
SPECTRAL CLASSIFICATIONS AND EXTINCTION DATA FOR POLARIMETRY PROGRAM STARS

Star	Spectral Type	V	E_{B-V}	E_{V-K}	E_{J-K}	A_V^a	R_V	References
HD 28170	A3 V	8.91	0.44	1.20	0.23	1.32	3.00	1
HD 28225	A3 III	7.78	0.41	1.11	0.20	1.22	2.98	1, 2
HD 28975	A4 III	9.02	0.60	1.59	0.29	1.75	2.92	1, 2, 3
HD 29333	A2 V	8.79	0.62	1.80	0.32	1.98	3.19	2, 4
HD 29647	B6-7 IVp	8.38	1.00	3.30	0.69	3.63	3.63	5
HD 29835	K2 III	8.70	0.38	1.04	0.15	1.14	3.01	2, 6
HD 30168	B8 V	7.70	0.45	1.07	0.19	1.18	2.62	1, 6
HD 30675	B3 V	7.52	0.51	1.40	0.25	1.54	3.02	1, 2
HD 279652	A2 V	9.82	0.31	0.90	0.17	0.99	3.19	7
HD 279658	B7 V	9.91	0.67	1.70	0.25	1.87	2.79	7
HD 283367	B9 V	10.34	0.65	1.77	0.37	1.95	3.00	2, 3
HD 283637	A0 V	11.25	0.78	2.07	0.39	2.28	2.92	2, 3
HD 283642	A3 V	10.60	0.68	1.99	0.37	2.19	3.22	6
HD 283643	A2 V	10.93	0.56	1.51	0.30	1.66	2.97	6, 8
HD 283701	B8 III	9.60	0.89	2.30	0.42	2.53	2.84	2
HD 283725	F5 III	9.96	0.61	1.59	0.30	1.75	2.87	1
HD 283757 ^b	A5 V	10.93	0.55	1.7:	...	6
HD 283800	B5 V	9.78	0.53	1.49	0.28	1.64	3.09	6
HD 283809	B3 V	10.78	1.61	5.18	1.00	5.70	3.54	9
HD 283812	A1 V	9.48	0.71	1.74	0.31	1.91	2.70	1, 2, 6
HD 283815	A5 V	9.75	0.57	1.74	0.30	1.91	3.35	6
HD 283855	A2	0.28	1.7:	...	2
HD 283877	F5 V	9.95	0.23	0.65	0.08	0.72	3.11	8
HD 283879	B5 V ^c	11.08	1.18	3.03	0.51	3.33	2.82	8, 10
Elias 3	K2 III	1.88	10.0:	...	11
Elias 29	G9 III	0.80	4.2:	...	11
Elias 30	K0 III	0.92	4.9:	...	11

^a Visual extinction is calculated from $A_V = 1.1E_{V-K}$ for all stars with both *UBV* and *JHK* photometry available. In other cases (values followed by colons): $A_V = 3.0E_{B-V}$ (stars lacking *JHK* photometry); $A_V = rE_{J-K}$ (stars lacking *UBV* photometry), with $r = 6$ ($E_{J-K} < 0.6$) and $r = 5.3$ ($E_{J-K} > 0.6$).

^b HD 283757 is the only star in our sample lacking infrared photometry: by chance it falls in a narrow strip of sky missed by 2MASS scans released to date.

^c No reliable spectral classification of HD 283879 exists in the literature: it is currently listed as an M star in Simbad, and as Ap, F5, or F8 by Slutsikii, Stal'bovskii, & Shevchenko (1980). Our adopted spectral type of B5 V is based on careful analysis of *UBV* and *JHK* colors. These colors are clearly incompatible with a classification as late as M for any value of reddening. The locus of the star in the $U-B$, $B-V$ plane is consistent with either \sim B5 or \sim F5; its locus in the $J-H$, $H-K$ plane is consistent with a reddened B-type star and inconsistent with an F-type star.

REFERENCES.—References for spectral types and *UBV* photometry are as follows: (1) Kenyon et al. 1994; (2) Simbad database; (3) Vrba & Rydgren 1985; (4) Oja 1987; (5) Crutcher 1985; (6) Straizys & Meistas 1980; (7) Ungerer et al. 1985; (8) Slutsikii et al. 1980; (9) Straizys et al. 1985; (10) This paper, see footnote c; (11) Elias 1978a.

absorption feature. Recent observations reported by Murakawa, Tamura, & Nagata (2000) have greatly increased the number of stars with $\tau_{3.0}$ and A_V values available, and a reassessment is therefore timely. Figure 5 plots $\tau_{3.0}$ versus A_V for all available lines of sight with $A_V > 2$ and $\tau_{3.0} \geq 0.05$, using data from Whittet et al. (1988), Smith et al. (1993) and Murakawa et al. (2000). The Smith et al. sample (nine stars) is a subset of the Whittet et al. sample (12 stars), and these have been combined (*filled circles*), taking average values of $\tau_{3.0}$ for common stars; A_V is taken from our current work as available (Table 4: HD 29647, HD 283809, Elias 3) or otherwise estimated from equation (8) with $r = 5.3$. The Murakawa et al. results (*open circles*) are presented as they appear in Table 2 of that paper, except that we have substituted our optically-determined A_V values for their values (based on $A_V = 12E_{H-K}$) for HD 29647 and HD 283809.

The distribution of points in Figure 5 shows no evidence for a systematic difference between the data sets that might arise, for example through different methods of evaluating A_V . The unweighted linear least squares fit to all 36 points,

shown in Figure 5, is represented by

$$\tau_{3.0} = q(A_V - A_{th}), \quad (9)$$

with $q = 0.072 \pm 0.002$ and $A_{th} = 3.2 \pm 0.1$. This result for the threshold extinction is in good agreement with the value of 3.3 originally proposed by Whittet et al. (1988) for a much smaller sample. There does not appear to be strong evidence in Figure 5 for intrinsic variations proposed by Murakawa et al. (2000) in the distribution of ice with A_V in the dark-cloud environment. We therefore adopt $A_{th} = 3.2$.

3.2. The Ratio of Total to Selective Extinction

R_V -values measured in dense clouds generally range from the diffuse-ISM average of about 3.0 up to \sim 5.5 in some regions of active star formation such as M42 (e.g., Whittet 1992 and references therein). Systematic increases in R_V with total extinction are predicted if stars with high A_V sample dense regions of the cloud, as the rate of grain growth (§ 1) should increase with density. Such behavior has been reported previously for other dark clouds, most

TABLE 5
EXTINCTION DATA FOR ADDITIONAL REDDENED STARS

Star	Spectral Type	V	E_{B-V}	E_{V-K}	E_{J-K}	A_V^a	R_V	References
HD 27405.....	B8 V	7.79	0.25	0.65	0.07	0.72	2.86	1
HD 27778A.....	B3 V	6.39	0.38	1.02	0.17	1.12	2.95	1
HD 27778B.....	A0 V	8.19	0.32	0.85	0.10	0.94	2.92	1
HD 27877.....	B9 V	7.36	0.42	1.27	0.24	1.40	3.33	1, 2, 6
HD 27923.....	B9 V	8.13	0.27	0.64	0.10	0.70	2.61	1
HD 28149.....	B5 V	5.59	0.11	0.24	0.04	0.26	2.4:	1
HD 28272.....	A6 V	9.02	0.39	0.98	0.20	1.08	2.76	1
HD 28482.....	B7 V	7.13	0.56	1.41	0.28	1.55	2.77	1
HD 28991.....	F3 V	7.95	0.05	0.14	0.02	0.15	3.1:	1
HD 30067.....	A2 V	8.97	0.36	0.87	0.15	0.96	2.66	1
HD 31252.....	A0 V	8.47	0.32	0.99	0.15	1.09	3.40	1, 2
HD 283550.....	A2 V	8.46	0.50	1.34	0.25	1.47	2.95	1
HD 283553.....	B6 V	9.07	0.41	1.01	0.14	1.11	2.71	1, 2
HD 283592.....	A4 V	9.86	0.26	0.74	0.12	0.81	3.13	1
HD 283740.....	A0 V	10.35	0.36	0.98	0.16	1.08	2.99	1
HD 283920.....	A0 V	9.12	0.61	1.49	0.20	1.64	2.69	1
HD 283929.....	A5 V	9.37	0.45	1.27	0.19	1.38	3.06	1, 2, 6
HD 283932.....	B7 V	9.92	0.75	2.04	0.32	2.24	2.99	1
HD 283939.....	A2 V	10.29	0.41	1.27	0.22	1.40	3.41	1, 2
HD 284012.....	B6 V	8.29	0.62	1.74	0.32	1.91	3.09	1
HD 284014.....	B5 V	9.35	0.79	2.00	0.41	2.20	2.78	1
HD 284287.....	F2 III	8.90	0.25	0.70	0.14	0.77	3.08	1
VR c.....	A0	10.18	1.00	2.79	0.53	3.07	3.07	3
VR d.....	A1	11.88	0.66	1.96	0.40	2.16	3.27	3
VR e.....	A7	11.48	0.68	1.78	0.34	1.96	2.88	3
091085.....	G0	16.96	1.54	4.64	0.97	5.10	3.31	12
141134.....	A5	14.29	1.19	4.15	0.95	4.57	3.84	12

^a Calculated from $A_V = 1.1E_{V-K}$.

REFERENCES.—(1) Kenyon et al. 1994; (2) Simbad database; (3) Vrba & Rydgren 1985; (4) Oja 1987; (5) Crutcher 1985; (6) Straizys & Meistas 1980; (7) Ungerer et al. 1985; (8) Slutskii et al. 1980; (9) Straizys et al. 1985; (10) This paper, see footnote c; (11) Elias 1978a; (12) Arce & Goodman 1999.

notably ρ Ophiuchi (Carrasco et al. 1973; Vrba, Croyne, & Tapia 1993), where $R_V \sim 4.0$ – 4.5 is typical of the densest regions. As the highest value found for any star in our sample is a mere 3.8, it is evident that the TDC is dominated by dust broadly similar to that in the diffuse ISM, at least

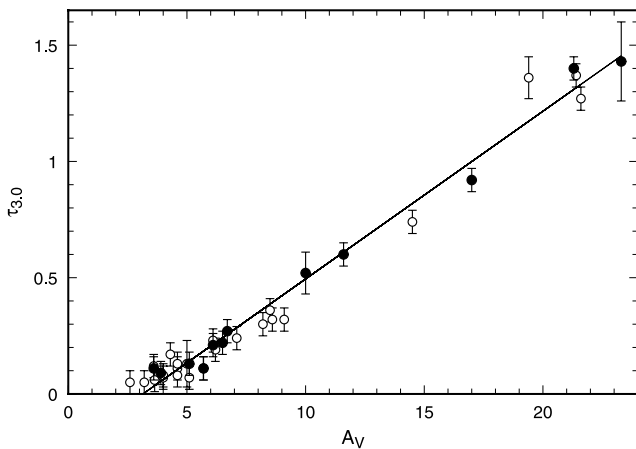


FIG. 5.—Plot of ice optical depth $\tau_{3.0}$ vs. extinction A_V for the Taurus dark cloud. *Filled circles*: 12 stars from the sample of Whittet et al. (1988), with $\tau_{3.0}$ values averaged from Whittet et al. (1988) and Smith et al. (1993); error bars represent scatter between the two sets of observations where this exceeds 0.05. *Open circles*: 24 stars from the sample of Murakawa et al. (2000). Only stars with $A_V > 2$ and $\tau_{3.0} \geq 0.05$ are included. The line is the unweighted linear least-squares fit to all data points. The intercept on the A_V axis yields a threshold extinction of $A_{th} = 3.2 \pm 0.1$.

over the extinction range sampled. In this section, we examine in detail the evidence for systematic changes in R_V .

The frequency distribution of R_V in the TDC is plotted as a bar chart in Figure 6. For comparison, we show the distribution for a typical sample of the diffuse-ISM (toward highly reddened OB stars; He et al. 1995), represented by a gaussian profile centered at $R_V = 3.05$ with $\sigma = 0.3$ (dashed

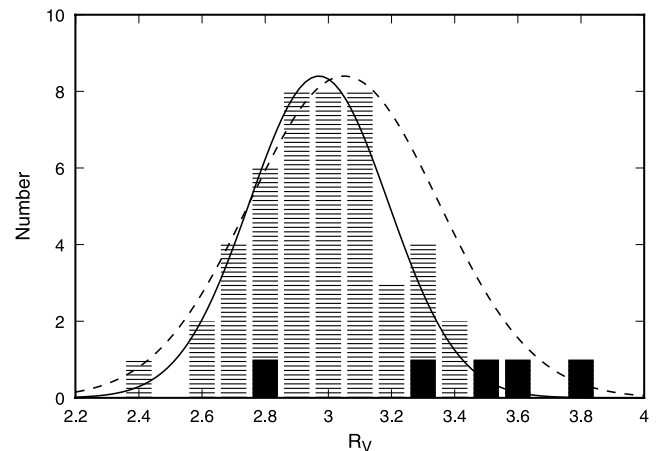


FIG. 6.—Bar chart of R_V values for stars in the Taurus cloud, obtained by plotting numbers of stars in bins of width 0.1. Stars with $A_V > 3.2$ (i.e., above the ice threshold) are shown as solid bars. The continuous curve is a Gaussian distribution (centered at $R_V = 2.97$ with $\sigma = 0.22$) fitted by eye to the Taurus data. The dashed curve (centered at $R_V = 3.05$ with $\sigma = 0.3$) represents the diffuse-ISM data of He et al. (1995).

curve). Although this curve fits reasonably well on the low- R_V side of the peak, the TDC data are better represented by a narrower profile ($\sigma = 0.22$) centered at 2.97 (continuous curve). Stars with $3.1 < R_V < 3.5$ seem to be somewhat under-represented in the TDC compared with the diffuse ISM. Note, however, that four out of five stars with $A_V > A_{th}$ have above-average R_V values.

The E_{B-V} versus E_{V-K} diagram, plotted in Figure 7, provides the clearest evidence for a real change in the optical properties of the dust as the extinction increases. The plot should resemble a straight line through the origin of slope $1.1/\langle R_V \rangle$ (by analogy with Fig. 4) if R_V is the same everywhere and scatter is due only to random observational error. The distribution of points in Figure 7 suggests that this prediction does, indeed, hold true for stars with $E_{V-K} < 2.9$ (equivalent to $A_V < 3.2$, i.e., below the ice threshold). The least-squares fit through the origin to these stars (solid line) yields $\langle R_V \rangle = 2.97 \pm 0.15$ (consistent with the peak in Fig. 6). At greater E_{V-K} , the slope appears to change systematically in the direction of higher R_V . Note that the four stars with highest E_{V-K} values in Figure 7 (E_{B-V} vs. E_{V-K}) are displaced (relative to the correlation line) in the opposite sense compared with their displacements in Figure 4 (E_{J-K} vs. E_{V-K}), i.e., increasing R_V appears to be accompanied by decreasing r .

The R_V values for individual stars are plotted against A_V in Figure 8. Points with $A_V < 3.2$ are distributed about $R_V = 2.97$, represented by the horizontal line in Figure 8. Scatter might arise in part from real star-to-star differences in addition to observational error, but no systematic trends are apparent for the low- A_V stars. The vertical line represents the ice threshold extinction. Two stars with $A_V \approx A_{th}$ have normal R_V , whilst four stars with $A_V > A_{th}$ have above-average R_V . A logical interpretation of the data is that only stars with extinctions appreciably above the ice threshold sample regions dense enough for systematic grain growth leading to enhanced R_V .

3.3. The Wavelength of Maximum Polarization

Previous work has shown that λ_{max} is generally quite well correlated with R_V , consistent with the expectation that each of these parameters varies in proportion to mean grain

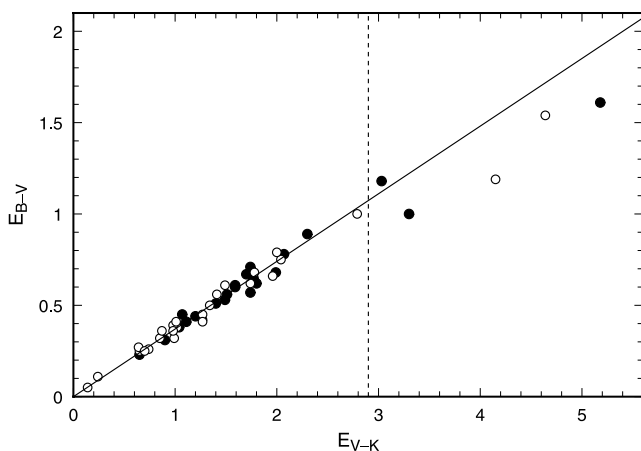


FIG. 7.—Plot of E_{B-V} vs. E_{V-K} . Plotting symbols have the same meaning as in Fig. 4. The vertical dashed line represents the ice threshold ($E_{V-K} = A_{th}/1.1 = 2.9$). The solid line is the linear least-squares fit through the origin for points below the ice threshold, with slope equivalent to $\langle R_V \rangle = 2.97 \pm 0.15$.

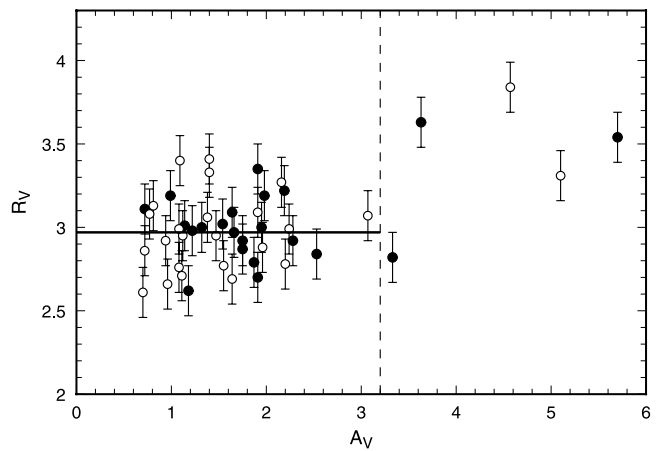


FIG. 8.—Plot of R_V vs. A_V . Symbols have the same meaning as in Fig. 4. The vertical dashed line represents the ice threshold extinction, $A_{th} = 3.2$. The horizontal line represents the mean value $\langle R_V \rangle = 2.97$ for $A_V < A_{th}$, deduced from Fig. 7.

size (Serkowski 1973; Serkowski et al. 1975; Whittet & van Breda 1978; Clayton & Mathis 1988; Vrba et al. 1993; Whittet et al. 1994; Larson, Whittet, & Hough 1996). Whittet & van Breda (1978) found that

$$R_V = (5.6 \pm 0.3)\lambda_{max} \quad (10)$$

adequately describes this correlation for a variety of interstellar environments from diffuse to dark clouds, although the scatter is often greater than can be explained purely in terms of observational error (see Clayton & Mathis 1988). The plot of R_V versus λ_{max} for the TDC is shown in Figure 9 (HD 279652 is omitted from this and subsequent plots because of the large uncertainty in λ_{max}). Remarkably, Figure 9 shows little correlation. There is a clustering of points around the ISM average ($R_V \approx 3.1$, $\lambda_{max} \approx 0.55 \mu\text{m}$), and the highest- R_V star (HD 29647) also has the highest λ_{max} ; but otherwise there is no clear trend of increasing R_V with λ_{max} (compare the predicted correlation from eq. [10]). Apparently, λ_{max} can have a wide range of values (~ 0.50 – $0.71 \mu\text{m}$) toward stars with R_V -values that do not deviate

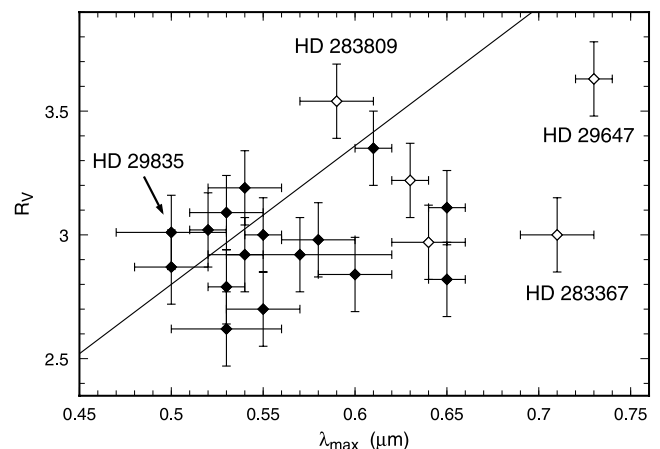


FIG. 9.—Plot of R_V vs. λ_{max} for stars in the polarimetry sample. Data are from Tables 4 and 2, respectively. Open symbols represent five stars with significant position-angle rotation with respect to wavelength (§ 2.1), filled symbols those with essentially constant position angle. The diagonal line is the mean relationship $R_V = 5.6\lambda_{max}$ deduced by Whittet & van Breda (1978).

significantly from the mean. Compare, for example, HD 283367 and HD 283725 in Figure 9: both have R_V values close to 3.0 but their λ_{\max} values differ by $\sim 40\%$.

We hypothesize that the unexpectedly poor degree of correlation between λ_{\max} and R_V is caused by *size-dependent variations in grain alignment*. In the past, λ_{\max} has been widely used as a proxy for mean grain size, but λ_{\max} is actually capable of two kinds of behavior, which we term “size-dependent” and “alignment-dependent” behavior. Previous discussions have emphasized size-dependent behavior. However, several studies have shown that the polarizing efficiency of interstellar dust, as measured by the ratio p/A at some wavelength, declines systematically with total extinction, as one probes progressively denser environments within a dark cloud (e.g., Goodman et al. 1992; Whittet et al. 1994; Gerakines, Whittet, & Lazarian 1995; Arce et al. 1998). That this holds true for the present sample is illustrated in Figure 10, which plots p_{\max}/A_V against A_V . The curve represents a fit of the form adopted by Gerakines et al. (1995). Note that for $A_V < A_{\text{th}}$, p_{\max}/A_V has a wide range of values, but for $A_V > A_{\text{th}}$, p_{\max}/A_V is systematically low ($< 1.5\% \text{ mag}^{-1}$; for comparison, the general interstellar mean is $\sim 3\% \text{ mag}^{-1}$). A number of factors may affect polarizing efficiency in a given line of sight. Stars in our sample with significant θ -rotation have low p_{\max}/A_V , suggesting depolarization resulting from superposition of distinct cloud components (Messinger, Whittet, & Roberge 1997). However, a systematic decline in polarization efficiency with increasing A_V is evident even if stars with θ -rotation are ignored. The most probable cause is a corresponding reduction in the efficiency of the alignment mechanism (see Gerakines et al. 1995 for discussion). Although the alignment mechanism is not fully understood (Roberge 1996), the degree of alignment achieved in a given environment is predicted to be a strong function of grain size for either superparamagnetic or suprathreshold alignment (Mathis 1986; Lazarian & Draine 1997), i.e., in marginal conditions, it is the small grains that lose their alignment first. As the small grains polarize most efficiently at shorter wavelengths, λ_{\max} will shift toward longer wavelengths as their orientations become randomized. This change will be independent of R_V if the size distribution of

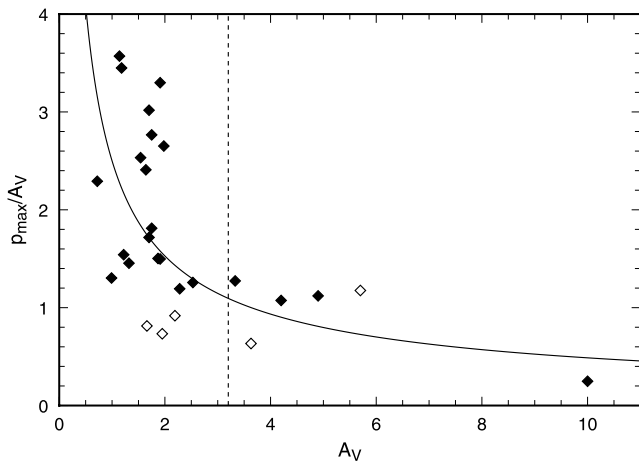


FIG. 10.—Plot of polarizing efficiency (p_{\max}/A_V) vs. A_V . Symbols have the same meaning as in Fig. 9. The dashed vertical line represents the ice threshold extinction, $A_{\text{th}} = 3.2$. The curve is a fit of the form $y = ax^b$ (Gerakines et al. 1995) with $a = 2.5$ and $b = -0.71$.

all grains (summed over aligned and unaligned components) remains the same.

Figures 11 and 12 compare correlations of λ_{\max} with extinction A_V and polarizing efficiency p_{\max}/A_V , respectively. The scatter in λ_{\max} above the value predicted by R_V (eq. [10]) at low A_V is evident in Figure 11 (compare Fig. 8). Whereas Figure 11 shows a weak positive correlation, Figure 12 shows a rather stronger *negative* correlation (the correlation coefficients are 0.45 and -0.76 , respectively). To test whether this situation is peculiar to the TDC, we examined corresponding plots for the Chamaeleon I dark cloud using data from Whittet et al. (1994), and found it to exhibit similar behavior: the correlation coefficients are 0.62 (λ_{\max} vs. A_V) and -0.67 (λ_{\max} vs. p_{\max}/A_V) for Chamaeleon I. These results are consistent with the hypothesis of high λ_{\max} arising from reduced alignment efficiency.

Does λ_{\max} carry any useful information on the *integrated* grain size distribution in the TDC that can be used to study grain growth? Our results clearly indicate that it cannot be

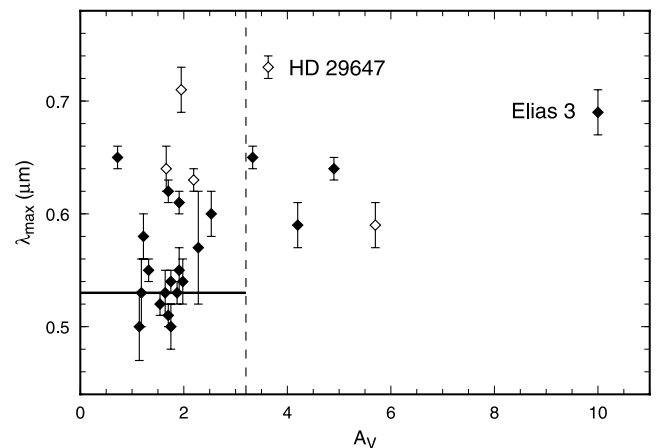


FIG. 11.—Plot of λ_{\max} vs. A_V . Symbols have the same meaning as in Fig. 9. The dashed vertical line represents the ice threshold extinction, $A_{\text{th}} = 3.2$. The horizontal line represents the value of λ_{\max} equivalent to the mean value of R_V using equation (10), i.e., $\langle \lambda_{\max} \rangle = 2.97/5.6 = 0.53 \mu\text{m}$.

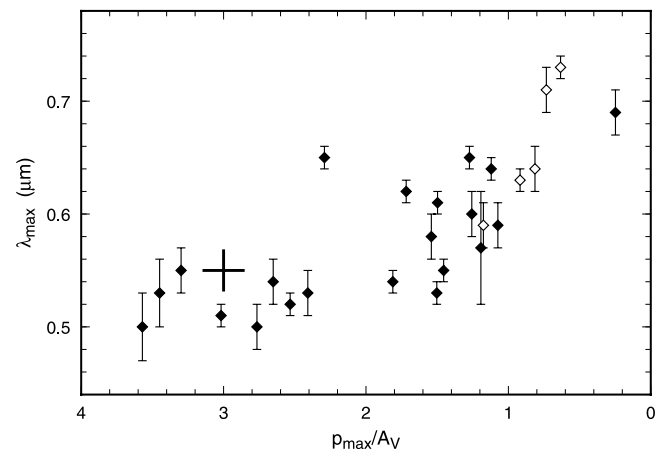


FIG. 12.—Plot of λ_{\max} vs. p_{\max}/A_V . The cross indicates the locus of the general diffuse-ISM average, other symbols have the same meaning as in Fig. 9. The p_{\max}/A_V axis is reversed to facilitate comparison with other plots, i.e., extinction increases from left to right.

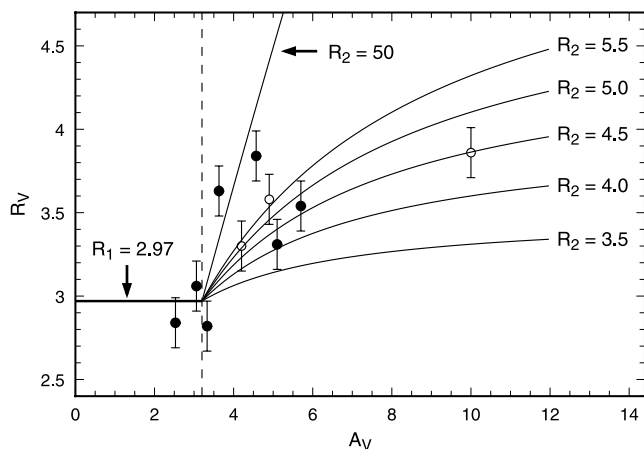


FIG. 13.—Dependence of R_V on A_V , according to the simple model described in § 4, for several values of the parameter R_2 . The dashed vertical line represents the ice threshold extinction, $A_{th} = 3.2$, and the horizontal line represents the mean value of R_V ($R_1 = 2.97$) at $A_V < A_{th}$. For clarity, only stars with $A_V > 2.5$ are plotted (filled circles: R_V determined photometrically; open circles: $R_V = 5.6\lambda_{max}$ assumed).

relied upon at extinctions $A_V < A_{th}$, where as many as 30%–40% of the stars have λ_{max} values significantly greater than predicted from their photometric R_V values. However, alignment-dependent behavior of λ_{max} appears to saturate quickly at higher A_V . Toward HD 29647 and HD 283809, enhanced λ_{max} is accompanied by enhanced R_V on approximately the expected scale (Fig. 9), indicating real changes in grain properties. The R_V value toward Elias 3 ($A_V \approx 10$) is unknown, but it is notable that its λ_{max} value, although high, is no greater than that of HD 29647 (see Fig. 11). An explanation for this might be that the alignment of grains deep within the cloud is so inefficient that dust at low-to-intermediate A_V dominates the polarization curve; in the case of Elias 3, this does seem probable given its low degree of polarization. However, mantled grains within the cloud must align to some extent, as they produce significant excess polarization corresponding to the $3 \mu\text{m}$ ice absorption feature (Hough et al. 1988; Kobayashi et al. 1999). With these caveats in mind, we tentatively assume that the $R_V \approx 5.6\lambda_{max}$ relationship gives a meaningful estimate of R_V for stars with $A_V > A_{th}$.

4. A SIMPLE MODEL FOR THE DEPENDENCE OF R_V ON A_V

A realistic model for the general structure of the TDC appears to be a screen of diffuse material in which clumps of denser material are embedded (Messinger et al. 1997; Arce & Goodman 1999; Murakawa et al. 2000). The observed dependence of polarization position-angle on wavelength toward HD 29647 and HD 283809 can be understood if both grain properties (size and/or composition) and the orientation of the magnetic field in the clump (TMC-1) differs from that in the screen (Messinger et al. 1997). The existence of an ice threshold extinction is consistent with this scenario: ice absorption is seen only in lines of sight that intercept a clump. Stars behind clumps have (at least) two components to their extinctions, one arising in the pervasive screen, and one in the clump (or clumps). We deduce the dependence of R_V on A_V , as predicted by this simple model, on the assumption that each component has

uniform optical properties. This assumption is undoubtedly simplistic, but it proves to be in reasonable accord with available observational constraints.

The ratio of total to selective extinction of a star behind a clump may be written

$$R_V = \frac{A_V}{A_1/R_1 + A_2/R_2}, \quad (11)$$

where $A_V = A_1 + A_2$ and subscripts 1 and 2 represent the diffuse screen and the clump, respectively. Note that “screen extinction” A_1 presumably arises in two parts, one in front of the clump and one behind, but this need not concern us here as the two parts add linearly. For $A_V > A_{th}$, we have $A_1 = A_{th}$ and $A_2 = A_V - A_{th}$, thus

$$R_V = \frac{A_V}{A_{th}/R_1 + (A_V - A_{th})/R_2}. \quad (12)$$

Empirical values for R_1 (2.97) and A_{th} (3.2) have been established by our work (§ 3.1–3.2). Equation (12) thus allows the dependence of R_V on A_V ($> A_{th}$) to be plotted for an assumed value of R_2 . For stars not behind clumps, $R_V = R_1 = 2.97$ ($0 < A_V < A_{th}$) is assumed.

Several values of R_2 were used in equation (12) to generate representative curves, plotted in Figure 13. Also shown for comparison is the distribution of empirical points for high-extinction stars ($A_V > 2.5$) in our sample. Figure 13 shows that the trend of increasing R_V predicted by the simple screen/clump model with $R_2 \sim 4.5$ is generally consistent with the observed behavior of R_V (or λ_{max}) with respect to A_V around or above the ice threshold. Two exceptions (HD 29647 and Arce-Goodman object 141134) require very much higher values of R_2 (~ 50) to give a sufficiently steep rise in R_V with A_V above A_{th} . It may well be the case that these lines of sight are not described adequately by our model. For example, the screen might be unusually thin, allowing a greater contribution to the total extinction from clump grains. A more intriguing possibility is that these objects might, in fact, be *embedded* within the cloud, in which case the effective value of A_{th} in equation (12) should be reduced (e.g., halved if equal columns of screen material lie in front of and behind the star); the curves in Figure 13 should then be displaced to the left, accordingly. However, at least in the case of HD 29647, this scenario seems inconsistent with available evidence, such as the lack of a reflection nebula (see Crutcher 1985 for discussion).

5. MANTLE GROWTH VERSUS COAGULATION: CLUES FROM UV EXTINCTION

Are the changes in mean grain size implied by increasing R_V toward stars with the highest dust-columns (Fig. 13) the result of mantle growth, coagulation, or a combination of both? It is certainly suggestive that R_V begins to change at $A_V \approx A_{th}$, the threshold extinction for the presence of H_2O -ice mantles. Thus, some fraction of the grains in the high- R_V lines of sight are evidently mantled. Extinction curves at ultraviolet (UV) wavelengths provide a valuable means of discrimination: the level of UV continuum extinction in normalized extinction curves should decline with increasing R_V if small grains are systematically incorporated into larger ones by coagulation (Cardelli et al. 1989; Mathis 1996), whilst the width of the 2175 \AA bump is sensitive to mantling (Mathis 1994). Regrettably, suitable UV spectra

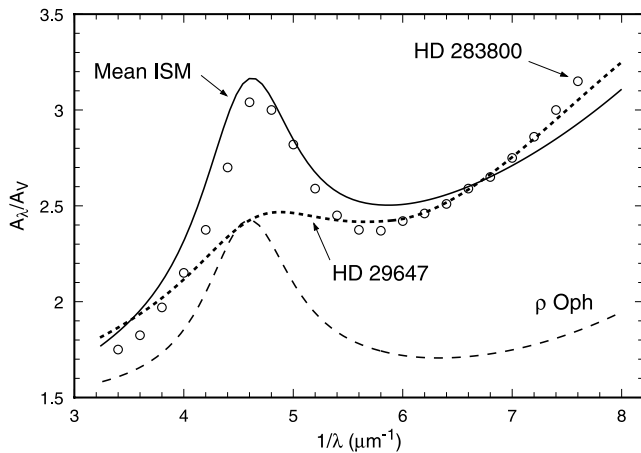


FIG. 14.— Ultraviolet extinction curves of Taurus stars HD 29647 and HD 283800, compared with those of HD 147933/4 (ρ Oph) and the mean interstellar medium. The HD 283800 curve is from Adamson (1999, private communication). The HD 29647, ρ Oph and mean ISM curves were calculated from empirical formulae (Fitzpatrick & Massa 1988), using parameters from Cardelli & Savage (1988) and Fitzpatrick & Massa (1988); this gives a close representation of the original data. Color excess ratios $E_{\lambda-V}/E_{B-V}$ were converted to normalized extinction A_λ/A_V using the appropriate R_V value in each case.

are available for only two lines of sight through the TDC: the extinction curve of HD 283800 was determined from spectra obtained with the *International Ultraviolet Explorer* satellite (*IUE*) by A. J. Adamson (1999, private communication), and that of HD 29647 (also determined from *IUE* data) was presented and discussed in detail by Cardelli & Savage (1988).⁹ Thus, we have one example of a star with $A_V < A_{th}$ and one with $A_V > A_{th}$, illustrated in Figure 14. Also shown are the mean diffuse-ISM and ρ Oph extinction curves (from Fitzpatrick & Massa 1988).

The extinction curve of HD 283800 (Fig. 14) is quite similar to the diffuse-ISM average curve, consistent with its normal R_V value (~ 3.1 ; Table 4), providing further illustration of the generally unremarkable properties of dust at $A_V < A_{th}$ in the TDC. In contrast, the extinction curves of ρ Oph and HD 29647 show substantial (and differing) departures from the norm:

ρ Oph.—The UV continuum extinction is relatively weak, whilst the 2175 Å bump has near-normal strength. The difference between the ρ Oph and mean-ISM curves is characteristic of the R_V -dependent behavior ($R_V \approx 4.1$ vs. 3.0) described by Cardelli et al. (1989). Coagulation is evidently the process that governs the grain size distribution toward ρ Oph (Jura 1980; Vrba et al. 1993), whereas ice-mantle growth appears to be unimportant at extinctions $A_V < 10$ (Harris, Woolf, & Rieke 1978; Tanaka et al. 1990). Coagulation naturally leads to reduced UV continuum extinction.

HD 29647.—The 2175 Å bump is much broader and weaker than in the mean ISM, whilst the UV continuum extinction is relatively strong. The breadth of the bump may

be attributed to mantling of small graphite grains,¹⁰ and the strength of the UV continuum extinction suggests that coagulation is relatively unimportant. These findings strengthen the argument that mantle growth is the dominant process in the TDC.

6. SUMMARY AND OVERVIEW

The results of this paper provide insight into the development of grain properties with extinction as we observe stars through progressively more opaque (and denser) regions of the TDC. At low A_V (0–3 mag), the dust is almost indistinguishable, in terms of its optical properties, from that in the general diffuse ISM. However, its alignment properties are already changing: this change is characterized by a systematic decline in polarizing efficiency with increasing A_V , accompanied by enhanced λ_{max} attributed to selective failure of the smaller grains to retain their alignment. For $A_V > 3$, real changes in optical properties of the grains occur, characterized by R_V values in the range 3.3–3.8. If the observed extinction toward each star includes a “normal” component as well as a “dense-clump” component, R_V may approach ~ 4.5 or more for dust within the clumps themselves. The available evidence favors mantle growth over coagulation as the dominant process affecting the optical properties of the dust in the TDC. If changes in R_V are not a direct consequence of mantle growth, then they must occur under closely similar physical conditions. Dust in Taurus appears to be in a different evolutionary state compared with other nearby dark clouds, such as ρ Oph, in which coagulation is the dominant process.

Despite our efforts to find information for highly reddened stars, our sample is heavily weighted toward stars with extinctions below A_{th} , which generally have low-to-normal R_V . It will be important in future work to determine reliable photometric R_V values for stars with higher extinctions ($A_V \gg A_{th}$), to better constrain the development of grain growth with physical conditions in the cloud. This will require accurate spectral classification and *BV* photometry for stars with $V \sim 16$. Another important goal for future studies will be to determine whether the UV extinction curves of stars with $A_V > A_{th}$ in the TDC show anomalous behavior similar to that seen so far only in HD 29647. HD 283809 is a prime candidate.

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⁹ Note that HD 29647 has a peculiar intrinsic spectrum with anomalously strong Hg and Mn lines (e.g., Crutcher 1985). This increases the problem of finding a suitable comparison spectrum to determine the UV extinction curve by the standard pair method, as discussed in detail by Cardelli & Savage (1988). These authors conclude that the UV extinction curve is truly anomalous and not significantly influenced by spectral mismatching.

¹⁰ Note that pure ice mantles cannot produce the observed increase in the width of the 2175 Å bump toward HD 29647: Mathis (1994) proposes that the required change in optical constants is caused by incorporation of neutral polycyclic aromatic hydrocarbons into the mantles.

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