



# IUPAC Critical Evaluation of the Rotational-Vibrational Spectra of Water Vapor. Part II. Energy Levels and Transition Wavenumbers for HD<sup>16</sup>O, HD<sup>17</sup>O, and HD<sup>18</sup>O

Jonathan Tennyson,<sup>a</sup> Peter F. Bernath<sup>b</sup>, Linda R. Brown<sup>c</sup>, Alain Campargue<sup>d</sup>, Attila G. Császár<sup>e</sup>, Ludovic Daumont<sup>f</sup>, Robert R. Gamache<sup>g</sup>, Joseph T. Hodges<sup>h</sup>, Olga V. Naumenko<sup>i</sup>, Oleg L. Polyansky<sup>j,a</sup>, Laurence S. Rothman<sup>k</sup>, Robert A. Toth<sup>c</sup>, Ann Carine Vandaele<sup>l</sup>, Nikolai F. Zobov<sup>j</sup>, Sophie Fally<sup>m</sup>, Alexander Z. Fazliev<sup>i</sup>, Tibor Furtenbacher<sup>e</sup>, Iouli E. Gordon<sup>k</sup>, Shui-Ming Hu<sup>n</sup>, Semen N. Mikhailenko<sup>i</sup>, Boris A. Voronin.<sup>i</sup>

<sup>a</sup>*Department of Physics and Astronomy, University College London, London WC1E 6BT, United Kingdom*

<sup>b</sup>*University of York, York, United Kingdom*

<sup>c</sup>*Jet Propulsion Laboratory, California Institute of Technology, Pasadena, CA, U.S.A.*

<sup>d</sup>*Université Joseph Fourier, Grenoble, France*

<sup>e</sup>*Loránd Eötvös University, Budapest, Hungary*

<sup>f</sup>*Université de Reims Champagne-Ardenne, Reims, France*

<sup>g</sup>*University of Massachusetts, Lowell, MA, U.S.A.*

<sup>h</sup>*National Institute of Standards and Technology, Gaithersburg, MD, U.S.A.*

<sup>i</sup>*V. E. Zuev Institute of Atmospheric Optics, Russian Academy of Sciences, 1, Academician Zuev square, 634021, Tomsk, Russia*

<sup>j</sup>*Institute of Applied Physics, Russian Academy of Sciences, Nizhny Novgorod, Russia*

<sup>k</sup>*Harvard-Smithsonian Center for Astrophysics, Cambridge, MA, U.S.A.*

<sup>l</sup>*Institut d'Aéronomie Spatiale de Belgique, Brussels, Belgium*

<sup>m</sup>*Université Libre de Bruxelles, Brussels, Belgium*

<sup>n</sup>*Laboratory of Bond-Selective Chemistry, University of Science and Technology of China, Hefei, 230026, China*

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## Abstract

This is the second of a series of articles reporting critically evaluated rotational-vibrational line positions, transition intensities, pressure dependences, and energy levels, with associated critically reviewed assignments and uncertainties, for all the main isotopologues of water. This article presents

energy levels and line positions of the following singly deuterated isotopologues of water: HD<sup>16</sup>O, HD<sup>17</sup>O, and HD<sup>18</sup>O. The MARVEL (Measured Active Rotational-Vibrational Energy Levels) procedure is used to determine the levels, the lines, and their self-consistent uncertainties for the spectral regions 0–22 708, 0–1 674, and 0–12 105 cm<sup>-1</sup> for HD<sup>16</sup>O, HD<sup>17</sup>O, and HD<sup>18</sup>O, respectively. For HD<sup>16</sup>O, 54 740 transitions were analyzed from 76 sources, the lines come from spectra recorded both at room temperature and from hot samples. These lines correspond to 36 690 distinct assignments and 8 818 energy levels. For HD<sup>17</sup>O, only 485 transitions could be analyzed from 3 sources; the lines correspond to 162 MARVEL energy levels. For HD<sup>18</sup>O, 8 729 transitions were analyzed from 11 sources and these lines correspond to 1 860 energy levels. The energy levels are checked against ones determined from accurate variational nuclear motion computations employing exact kinetic energy operators. This comparison shows that the measured transitions account for about 86 % of the anticipated absorbance of HD<sup>16</sup>O at 296 K and that the transitions predicted by the MARVEL energy levels account for essentially all the remaining absorbance. The extensive list of MARVEL lines and levels obtained are given in the Supplementary Material of this article, as well as in a distributed information system applied to water, W@DIS, where they can easily be retrieved. In addition, the transition and energy level information for H<sub>2</sub><sup>17</sup>O and H<sub>2</sub><sup>18</sup>O, given in the first paper of this series [J. Quant. Spectr. Rad. Transfer 110 (2009) 573-596], has been updated.

*Key words:* Water vapor, transition wavenumbers, atmospheric physics, energy levels, MARVEL, information system, database, W@DIS, infrared spectra, microwave spectra, HD<sup>16</sup>O, HD<sup>17</sup>O, HD<sup>18</sup>O

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## 1. Introduction

The water molecule is the most abundant polyatomic molecule in the universe and the single most important species for controlling the Earth's climate. Thus, the spectrum of water is both one of the most important and one of the most thoroughly studied.

The first 14 authors of this article form a Task Group under the auspices of the International Union of Pure and Applied Chemistry (IUPAC) with the aim of constructing a database of water transitions from experiment and theory, with individual tasks described in Table 1 of the first paper in this series [1], henceforth known as Part I. Given the nature of water spectroscopy [2], this database will concentrate on the pure rotational and vibration-rotation transitions of water which we consider simultaneously on an equal basis. This paper is the second in a series presenting our evolving methods for collecting and analyzing the experimental and quantum chemical spectroscopic information as well as our validated data advocated for deposition in databases. In Part I, we derived energy levels and transition wavenumbers for the water isotopologues  $\text{H}_2^{17}\text{O}$  and  $\text{H}_2^{18}\text{O}$ . This was done using the Measured Active Rotational-Vibrational Energy Levels (MARVEL) protocol of Furtenbacher *et al.* [3, 4, 5], which was refined during Part I to allow for the treatment of larger datasets and to perform a significant amount of checking in order to minimize errors and inconsistencies in the initial experimental transition data.

In this work we apply MARVEL to the  $\text{HD}^{16}\text{O}$ ,  $\text{HD}^{17}\text{O}$ , and  $\text{HD}^{18}\text{O}$  isotopologues of the water molecule.  $\text{HD}^{16}\text{O}$ , singly deuterated water, has a fractional abundance in the earth's atmosphere, compared to  $\text{H}_2^{16}\text{O}$ , of about 0.0003 (note that the natural  $\text{HD}^{17}\text{O}/\text{H}_2^{17}\text{O}$  and  $\text{HD}^{18}\text{O}/\text{H}_2^{18}\text{O}$  abundances have a similar, about  $3 \times 10^{-4}$  value). Even though a trace species, atmospheric absorption by HDO can be significant. This is particularly true for  $\text{HD}^{16}\text{O}$  since some of its absorption spectrum, unlike that of  $\text{H}_2^{17}\text{O}$  and  $\text{H}_2^{18}\text{O}$ , is significantly shifted from that of  $\text{H}_2^{16}\text{O}$ . Observation of atmospheric

HDO spectra has long been used on both Earth [6] and other planets [7, 8] as a proxy to understand their climatic evolution. Atmospheric HDO spectra are also being extensively observed because of the information they provide on the transport of water vapor into the stratosphere [9, 10]. Abundance of HDO relative to H<sub>2</sub>O can vary significantly; the proportion of HDO on Venus, for example, is over one hundred times that of the terrestrial value [11].

Astrophysical observations of HDO are also potentially important.[9] Deuterium was formed in the first few minutes after the Big Bang. As D is burnt rapidly in stellar interiors, the abundance of D has been decreasing since this time. Deuterium burns in objects with masses greater than about 13 times that of Jupiter, which is one way of defining the boundary between a brown dwarf and a planet. Observations of HDO in the atmospheres of such objects, which are cool enough to be largely molecular but hot by terrestrial standards, could be the key to a deuterium test [12] that would distinguish between planets and other sub-stellar objects. In the cold interstellar medium fractionation effects become important leading to an expected overabundance of HDO [13]. Astronomical applications of water spectra therefore require spectral data appropriate for very low temperatures up to stellar temperatures of about 3000 K.

As emphasized in Part I, a distinguishing feature of the present series of IUPAC-sponsored studies is the joint utilization of all available experimental and the best theoretical line and level data, with a long-term aim to create a complete linelist for all water isotopologues. While determination of a complete linelist is outside the scope of present-day experiments, it can be determined by means of sophisticated first-principles quantum mechanical computations. Consequently, as long as experiments have a higher precision than even the most advanced computations that can be performed for a molecule of the size of water, the complete linelist will necessarily contain accurate experimental data and less accurate computational data. MARVEL-

type efforts help to (a) replace as many computed lines as possible with their experimental counterparts, (b) validate and ideally reduce the uncertainty with which a transition has been determined, and (c) facilitate the assignment of experimental spectra.

## 2. Methods, input data, and data treatment

The methods employed in this study for collecting and critically evaluating experimental transition wavenumbers and their uncertainties and for inverting the wavenumbers in order to obtain the best possible energy levels with uncertainties are based on the MARVEL procedure [3, 4, 5] involving the iterative robust reweighting scheme [14]. During a MARVEL analysis we simultaneously process *all* the available assigned experimental lines and the associated energy levels for the chosen isotopologue. MARVEL is not designed to determine any new energy levels. However, from the MARVEL energy levels obtained one can determine transitions whether they have been measured or not. The reweighting scheme means that uncertainties for certain selected transitions are changed (increased) during iterations of the MARVEL procedure. After cleansing of the database and applying the iterative robust reweighting algorithm of MARVEL, a database is created containing self-consistent and correctly assigned transitions and the seemingly best possible related uncertainties supported by the database. Energy levels, and their uncertainties, determined from these transitions are in harmony with the measured transitions and their (adjusted) uncertainties.

The first step in the MARVEL procedure is to split the transition data into spectroscopic networks (SN). SNs contain all interconnecting rotational-vibrational energy levels supported by the grand database of the transitions. For HDO, as there is no nuclear spin symmetry, the vast majority of the data forms a single SN. Other inter-connecting but unattached transition networks are designated as floating networks (FSNs) or, in the case of a single transition with no energy level in common with any of the other transitions

in the compilation, orphans (ORPs).

For HD<sup>16</sup>O there is a considerable number of experimental data sources available [15, 16, 17, 18, 19, 20, 21, 22, 23, 24, 25, 26, 27, 28, 29, 30, 31, 32, 33, 34, 35, 36, 37, 38, 39, 40, 41, 42, 43, 44, 45, 46, 47, 48, 49, 50, 51, 52, 53, 54, 55, 56, 57, 58, 59, 60, 61, 62, 63, 64, 65, 66, 67, 68, 69, 70, 71, 72, 73, 74, 75, 76, 77, 78, 79, 80, 81, 82, 83, 84, 85, 86, 87, 88, 89, 90, 91, 92, 93, 94, 95, 96, 97, 98, 99, 100, 101, 102, 103, 104, 105, 106, 107, 108, 109]. In particular, we note that a number of hot HD<sup>16</sup>O spectra [64, 82, 86, 106] are included in the above list. Hot spectra are usually considerably richer in transitions but have significantly larger uncertainties and a higher chance of misassignment than spectra recorded at room temperature. As expected, there is a much smaller number of publications reporting measured and assigned transitions for HD<sup>17</sup>O [61, 110, 111] and HD<sup>18</sup>O [42, 51, 61, 102, 108, 110, 112, 113, 114, 115].

Tables 2–4 provide, for each data source, experimental information related to the spectra of HD<sup>16</sup>O, HD<sup>17</sup>O, and HD<sup>18</sup>O, respectively. The number of originally measured (A) and validated (V) transitions for each data source is given there, as well. Due to the large number of related experimental studies, a nearly continuous coverage has been achieved for HD<sup>16</sup>O up to 22 800 cm<sup>-1</sup> but not for the other two isotopologues. As modern measurements of pure rotational transitions at microwave frequencies (< 30 cm<sup>-1</sup>) are very scarce, the microwave transitions were mostly obtained from early works and are limited in number. Conversely, early infrared measurements, such as that by Benedict *et al.* [29], have largely been superceded by more recent and more accurate measurements and were therefore largely omitted from the final compilation.

To be included in our tabulation, data sources must include original experimental line positions with uncertainties, line assignment, and information on the experimental conditions under which the experimental data were recorded. The latter data are summarized in the column ‘Physical condi-

tions' in Tables 2–4. In each case the data source is identified with a tag, as specified in Part I, based on year of publication and the names of the authors.

Abundance is a particular issue with spectra of HDO. Since a pure sample of HDO would partially disproportionate to  $\text{H}_2\text{O}$  and  $\text{D}_2\text{O}$ , it is not possible to record spectra of pure HDO. In practice, 50 % HDO is normally the best that can be achieved. The presence of significant quantities of  $\text{H}_2\text{O}$  and  $\text{D}_2\text{O}$  in any HDO sample means that a good knowledge of spectra of these species is a prerequisite for analysing spectra of HDO. Furthermore, it is very difficult to get complete spectral coverage for HDO since some transitions are obscured by strong lines associated with  $\text{H}_2\text{O}$  or  $\text{D}_2\text{O}$ .

Most of the spectra were obtained by Fourier transform spectroscopy (FTS), which gives a wide spectral coverage from the microwave region to the near ultraviolet. In order to detect weak lines, FTS spectrometers have been equipped with long multipass cells. Absorption path lengths as large as 433 and 1804 m have been achieved with the cells available at Kitt Peak and Reims, respectively, providing a large number of observed transitions, mostly in the near infrared (NIR) region.

Laser-based methods, such as CRDS (cavity ringdown spectroscopy) and ICLAS (intracavity laser absorption spectroscopy), are limited to certain spectral regions depending on the availability of tunable laser sources. These techniques have specific advantages in terms of sensitivity and spectral resolution, which make them particularly suitable for the characterization of spectral regions with weak absorption features. This is why extensive investigations with laser-based methods were mostly limited to transparency windows or to the visible region. In spite of the small natural abundance of even  $\text{HD}^{16}\text{O}$ , part of the data were obtained from spectra recorded with water in natural abundance.

### *2.1. Pre-MARVEL validation*

As for Part I, the experimental databases assembled for the three HDO isotopologues were checked for transcription problems or problems charac-

terizing the original data source. The following checks were made apart from the trivial checks for formatting incompatibilities and looking for entries with zero uncertainties. The dataset of transitions was searched looking for cases where the  $K_a + K_c$  sum, where  $K_a$  and  $K_c$  are the usual asymmetric top rigid rotor quantum numbers, had a value different from  $J$  or  $J + 1$ , where  $J$  is the rotational quantum number. We also searched for obvious duplications in the dataset and a warning was issued if this happened. Overall, fewer problems were detected for the HDO isotopologues than observed for  $\text{H}_2^{17}\text{O}$  and  $\text{H}_2^{18}\text{O}$  in Part I [1].

## 2.2. Assignment, labels

It is a requirement of the MARVEL protocol that the dataset contains a single unique assignment to label both the lower and the upper states involved in each transition. In this work we retain the order of vibrational labels applicable to  $\text{H}_2^{16}\text{O}$  for HDO, *i.e.*,  $v_1$ ,  $v_2$ , and  $v_3$  stand for the OD stretching, bending, and OH stretching quantum numbers, respectively, and they provide the vibrational label  $(v_1 v_2 v_3)$ . Note that in some works, such as Ref. [116],  $v_1$  and  $v_3$  are reversed. We use the standard asymmetric top quantum numbers  $[J K_a K_c]$  or  $J_{K_a K_c}$  to label the rotational states. Thus, the rotation-vibration levels of each isotopologue are identified uniquely by six labels altogether.

Before processing the published transition data we checked, as thoroughly as possible, whether the assignments were correct and consistent. Rovibrational labels, which can be used for checking the  $(v_1 v_2 v_3) [J K_a K_c]$  labels, can be taken from computations based on the use of an effective Hamiltonian (EH). In the EH approach, all vibrational states are combined in the polyads of interacting states based on the ratio between the harmonic frequencies  $\omega_1$ ,  $\omega_2$ , and  $\omega_3$ . For  $\text{HD}^{16}\text{O}$ ,  $\omega_1$  is at  $2823 \text{ cm}^{-1}$ , very close to twice  $\omega_2$ , which is at  $1444 \text{ cm}^{-1}$ , while  $\omega_3$ , at  $3887 \text{ cm}^{-1}$ , lies much higher. This structure leads to a series of well isolated  $(0 0 n v_3)$  vibrational states, whose energy levels can be easily fitted, and, consequently, labeled within the EH



approach. At the same time, the coincidence of  $\omega_1$  and  $2\omega_2$ , and the strong centrifugal distortion effects inherent in HDO, result in strong and numerous anharmonic resonances. These result in the formation of the resonance polyads which often include highly excited bending states. This complicates proper EH analysis and therefore inhibits the process of labeling for many of the high-lying energy levels. We note that to run MARVEL successfully such labels simply have to be self-consistent; the MARVEL process itself says nothing about the physical basis of the labels [117].

Validation of the assignments attached to the observed transitions was performed as follows. All transitions were examined for consistency of the upper levels derived from combination difference (CD) relations. This method is a simple and powerful tool for the assignment of rovibrational spectra; however, it is often inapplicable to weak transitions because of the incompleteness of the set of observed transitions. All the transitions associated with a given rotational level of the (0 0 0) vibrational ground state have been considered for combination differences. At this stage, conflicting labels could be easily traced and corrected. This step was helped considerably by the availability of results from variational nuclear motion computations [118]. Where necessary, labels were changed assuming similar increases in rotational energies as a function  $J$  and  $K_a$  for similar vibrational states with the same  $v_2$  quantum number as well as quasi-degeneracy of rotational levels with  $K_a$  close to  $J$  or  $K_a$  equal to 0 or 1. In the end, consistent labeling has been established for all the assignments considered. We recommend that the labeling provided in this paper should be generally adopted, although in cases of strongly perturbed energy levels there is some remaining uncertainty.

### 2.3. *Uncertainties*

Within the MARVEL protocol reasonable estimates for the accuracy of the observed transitions must be provided. Despite the adjustments by the robust reweighting scheme, false uncertainties attached to the transitions can noticeably deteriorate the accuracy of the MARVEL energy levels.

In the majority of the data sources proper experimental uncertainties are not given for each transition. Often only the general accuracy of the data for the region investigated is provided. For a few publications we were forced to estimate the experimental uncertainties. As no values were presented in the original source, these were based on average values characteristic of the experimental setup exploited in the measurement.

For further important adjustment of the uncertainties of the transitions see sections 2.6 and 2.8.

#### 2.4. Variational validation

As an independent validation of the experimental transition wavenumbers and the derived MARVEL energy levels, systematic comparisons were made with the results of state-of-the-art variational nuclear motion calculations. For this comparison the newly computed VTT HD<sup>16</sup>O linelist [118] was used; this linelist was computed using the spectroscopically determined HDO potential energy hypersurfaces of Yurchenko *et al.* [119], the so-called CVR dipole moment surface [120], and the DVR3D nuclear motion program suite [121].

When variational results are used for validation, we can rely on a well-known feature of such calculations: the smooth and slow variation of obs – calc residuals for the energy levels of a particular vibrational state having the same  $K_a$  and increasing  $J$  values [122]. The longest obs – calc sequences could be investigated for the hot spectra, where transitions involving  $J$  as high as 30 have been detected. Examples of the obs – calc deviations for the energy levels of the (0 0 0) state with  $K_a = 1 - 5$  are shown in Fig. 1. The obs – calc residuals for levels with a given  $K_a$  but different  $K_c$  increase as  $J$  increases, hindering the assignment of the dense observed spectrum without detailed consideration of these near degeneracies.

The trends in obs – calc for the highly excited vibrational states are not particularly smooth as they can be strongly perturbed by nearby states.

Examples of perturbations of the obs – calc trend due to resonance interactions can be observed in Fig. 2. Cases with erratic obs – calc trends were additionally checked to see whether the calculated energy level set includes the resonance partner, whose energy level has to be close to the level under investigation and whose quantum numbers should satisfy the conventional Coriolis-, Fermi-, or Darling–Dennyson-type resonance interaction rules, or some combination of them. We stress that so-called HEL (highly excited local) resonances [123], which do not obey conventional resonance selection rules and can involve energy levels with a large difference in the  $\nu_2$  bending quantum number, are especially strong in HDO due to the close coincidence between the  $\omega_1$  and  $2\omega_2$  harmonic frequencies [100].

Those transitions which involved a MARVEL energy level that did not have a matching variational counterpart within the obs – calc trend established for the corresponding vibrational state were “deleted” from the input to MARVEL (by adding a minus sign in front of the wavenumber) and the MARVEL process was repeated until all MARVEL levels had variational counterparts within the appropriate obs – calc trend; possible resonance distortions of these trends were also taken into account. The final, accepted obs – calc values vary in magnitude from 0.01 up to 0.32  $\text{cm}^{-1}$  depending on the vibrational state and the  $J$  and  $K_a$  values. A comparison of the experimental (MARVEL) and variational [118] energy levels is given in Fig. 3. For transitions removed at this stage, see the appropriate comments in the footnotes supplementing Tables 2–4. Levels which the Task Group considers to be marginally reliable were retained but are flagged in the appropriate entry (see the Supplementary Material).

### 2.5. Hot transitions

Unsurprisingly, the hardest levels to validate came from the hot transitions observed and assigned in 03JaTeBeZo [86], due to the high density of both predicted and observed transitions and the possibility of transitions to a given upper level originating from lower (sometimes unknown) levels belonging to

different vibrational states. During the process of comparing the observed and calculated energy levels it was discovered that 03JaTeBeZo often gave only one assignment to transitions involving quasi-degenerate lower and upper levels. This impeded proper evaluation of the second component of the quasi-degenerate pair and, consequently, led to the disconnection of many experimental energy levels which are connected to these second components. To solve this problem, the second components of the quasi-degenerate transitions omitted in 03JaTeBeZo were added to the data based on predictions of the VTT variational linelist. Experimental transitions were considered to be degenerate if their variational centers coincided within  $0.015 \text{ cm}^{-1}$ , since the smallest difference between experimentally observed lines in 03JaTeBeZo was always at least  $0.02 \text{ cm}^{-1}$ . By this means about 1600 extra line assignments were added to the 03JaTeBeZo list (marked with D, see the Supplementary Material).

For the validation procedure we also used the unpublished experimental linelist of hot HDO emission spectra recorded by Mellau in the  $400\text{--}900 \text{ cm}^{-1}$  region [64]. Due to peculiarities of the hot spectrum assignment, deleting one unconfirmed experimental transition may lead to the loss of many experimental energy levels derived in 03JaTeBeZo. For this reason we assigned 65 transitions using the unassigned lines from the experimental linelist attached to 03JaTeBeZo (for details see the Supplementary Material).

### 2.6. Recalibration

As discussed in Part I, the absolute accuracies of different datasets vary, producing systematic offsets in the observed line positions. In massive compilations to gather measurements obtained over decades, these differences can arise partly because calibration standards improve over time. In the early 1980's, advanced techniques produced frequency measurements that became the dominant standards for infrared spectroscopy with reported accuracies between  $3 \times 10^{-7}$  and  $2 \times 10^{-5} \text{ cm}^{-1}$ :  $\text{CO}_2$  at  $10 \mu\text{m}$  [124], the P(7) line of  $\text{CH}_4$  at  $3.34 \mu\text{m}$  [125], and the  $2-0$  band of  $\text{CO}$  at  $2.3 \mu\text{m}$  [126]. Using these

frequency standards, transitions of other gases were recalibrated but often with lesser precision ( $1 \times 10^{-4} \text{ cm}^{-1}$ ). With FTS data, such corrections are easily made just by applying a multiplicative factor obtained as a ratio of the new to the old positions. Water, being ubiquitous in spectrometer chambers and having many strong bands throughout the infrared, was reinvestigated against the then-exceptional quality frequencies [124, 125, 126]. For example, the best prior  $\nu_2$  band line positions near  $6 \mu\text{m}$  [47] were altered by a factor of 0.999 999 77 [127]. Later, newer technological developments resulted in calibration standards in the near-infrared, such as  $\text{C}_2\text{H}_2$  near  $1.5 \mu\text{m}$  [128] and atomic potassium near  $0.77 \mu\text{m}$  [129].

During the course of this study it became apparent that there are several sources of HDO data which might suffer from calibration problems. The sources identified include 78KaKaKy [43], 83Guelachv [47], 93Tothb [61], 86FlCaMaGu [52], 89OhSa [55], 91RiSmDeBe [58], 99HuLiHeCh [70], and 99Toth [72]. It is straightforward to determine multiplicative calibration factors with MARVEL by minimizing the root-mean-square (rms) deviation between the observed transitions with wavenumbers scaled with a given calibration factor and those produced by MARVEL from the energy levels. This minimization was performed sequentially for all the problematic data sources; Table 1 gives the calibration factors we determined for cases where our analysis gave a result which differed significantly from unity. For the sources 91RiSmDeBe, 93Tothb, and 99Toth the MARVEL analysis suggests that no recalibration is needed or possible; thus, these transitions were not altered. For 78KaKaKy, 83Guelachv, 86FlCaMaGu, and 89OhSa the experimental FTS data were adjusted using the calibration factors determined. Only the recalibrated results were used in the final data analysis.

Improvement of the ICLAS data via recalibration was attempted only in one case here [75]. This is despite the fact that there are known calibration problems with some of the ICLAS data [100]. These arise because different calibration lines are used for every few  $\text{cm}^{-1}$ , making it difficult to determine

Table 1: Experimental sources recalibrated for the three isotopologues of HDO during the course of this work

Tag	Range ( $\text{cm}^{-1}$ )	Calibration factor	Shift <sup>a</sup> ( $\text{cm}^{-1}$ ) ( $\text{cm}^{-1}$ )	Median <sup>c</sup>	Original standard
78KaKaKy [43]	151–420	1.000 005 413	+0.0013	293	
86FlCaMaGu [52]	1125–1890	0.999 999 77	−0.00035	29	[47]
91RiSmDeBe [58]	1176–1650	0.999 999 98	−0.00003 <sup>b</sup>	49	[132]
83Guelachv [47]	1191–1740	0.999 999 77 <sup>d</sup>	−0.00035	45	[133]
89OhSa [55]	6384–6596	0.999 998 93	−0.0069	23	[134]
99HuLiHeCh[70]	12525–12912	1.000 000 819 <sup>e</sup>	+0.0104	7	[135]
00NaBeCaSc[75]	–	0.999 999 452 <sup>f</sup>			

<sup>a</sup> Shift of the middle wavenumber value of the range (new – old).

<sup>b</sup> This shift is within the original precision of about  $0.00006 \text{ cm}^{-1}$ .

<sup>c</sup> Indicates that half of the energy levels involved in the transitions in the given source are part of at least this many measured and validated transitions.

<sup>d</sup> The same calibration factor was determined in [130].

<sup>e</sup> A very similar calibration factor was determined in [100]. <sup>f</sup> This is the only ICLAS measurement which was recalibrated as part of the present study.

a unique calibration factor for the whole region covered or employ a constant shift value.

From time to time, the available frequency standards are reevaluated producing new recommendations (e. g., [130, 131]). The last composite review [130] was performed 15 years ago. Thus, we must be mindful that ongoing advancements in frequency measurements may later reveal the need for further recalibration of the HDO data collected here.

### 2.7. *Post-MARVEL validation*

While data handling within the MARVEL process is performed as automatically as possible, for a number of transitions, which proved to be clear outliers by combination difference relations, the experimental uncertainties were increased manually (see footnotes to Tables 2–4) by a process we call post-MARVEL validation.

This extra validation and the subsequent adjustment was done when the energy of an upper rotational-vibrational state deviated far more from the corresponding mean value established by the lower MARVEL energy levels plus the transition wavenumbers than the stated experimental uncertainty. The MARVEL protocol can make similar adjustments automatically if the error associated with a transition is an outlier and all the data have similar accuracy. However, a problem arises if an erroneously small experimental uncertainty is attached to what is actually a much less accurate experimental datum, the same level is involved in several measurements, and other transitions in the combination difference relations, though are consistent, have formally much larger uncertainties. In this case the MARVEL energy level will be determined by the formally most accurate transition which, in fact, represents an outlier.

For HDO, all data declared to have very high accuracy, such as those from 07VoNaCaCo [100], 07JeDaReTy [101], 97Tohc [67], and 07MiWaKaCa [103], were subject to post-MARVEL checks. In this context we note that sometimes the experimental uncertainty attached to a line in the original source reflects the quality of the line profile fit rather than the real accuracy with which the wavenumber was determined.

This validation was not performed in Part I. Thus, in the active update of the data of  $\text{H}_2^{17}\text{O}$  and  $\text{H}_2^{18}\text{O}$  presented in the Appendix of this paper the energy levels were subjected to a post-MARVEL validation.

Table 2: Data sources and their characteristics for HD<sup>16</sup>O. See section 2.8 for comments.<sup>a</sup>

Tag	Range (cm <sup>-1</sup> )	Trans. A / V	Physical conditions				Comments
			T(K)	p(hPa)	Abun.	Rec.	
62TrBe [32]	0.016–0.028	2 / 2	RT			SMM	(2a)
55WeBeHe [27]	0.016–0.102	6 / 1	RT			SMM	(2b)
53BeWe [21]	0.027–0.101	5 / 4	RT			SMM	(2c)
54PoSt [26]	0.19–0.90	7 / 7	RT			MS	(2d)
68VeBI Dy [35]	0.286	1 / 1	RT			BMS	(2e)
53Crawford [24]	0.295	3 / 1	RT			SMM	(2f)
57Posener [30]	0.343	1 / 1	RT			SMM	(2g)
64ThKrLo [33]	0.343	1 / 1	RT			BMS	(2h)
67BI Ve Dy [34]	0.343	1 / 1	RT			BMS	(2i)
49Strandbe [18]	0.343–1.676	3 / 3	RT	29.33		SMM	(2j)
71LaBeSt [39]	0.35–2.24	9 / 9	RT				
70BeSt [37]	0.682–5.344	3 / 3	RT			MS	
46ToMe [15]	0.744	1 / 1	RT				(2k)
53JeBiMa [22]	0.830–0.89	2 / 2	RT		50%	MS	(2l)
49Mcafee [17]	0.90–0.90	1 / 1					
56ErCo [28]	1.674–8.052	6 / 6	RT			MS	
70StBe [36]	1.730–8.872	11 / 11	RT			MS	
73ClBeKlRo [40]	2.93–8.872	6 / 0	RT			SMM	
71DeCoHeGo [38]	2.93–25.13	36 / 36	RT		50%	TMS	
93GoFeDeDu [62]	7.54–8.07	3 / 3	100-600	0.013-1.3		CCC	(2m)
87BaAlAlPo [54]	7.680–11.527	3 / 1	1000			MW	(2n)
84MeDeHe [49]	0.39–33.67	32 / 32				MW	(2o)
76FlGi [42]	15.51–39.37	20 / 13	RT	13.3	50%	FS	(2p)
85Johns [51]	20–339	225 / 225		1.33		FTS	(2q)
78KaKaKy [43]	152–419	60 / 60	RT	0.53-5.3		FTS	0.50 (2r)
95PaHo [63]	171–498	335 / 335	295	0.5	50%	FTS	(2s)
03JaTeBeZo [86]	381–3932	11503/11184	1770	3.33	25%	FTS	(2t)
99Toth [72]	657–2066	749 / 748	296	0.6-6.8	40-47%	FTS	(2u)
93Tothb [61]	969–1892	1267/1265	297	0.4-21.2	10-49%	FTS	0.25-2.39 (2v)
86FlCaMaGu [52]	1125–1890	1253/1252	RT	7.6		FTS	(2w)
91RiSmDeBe [58]	1176-1650	266 / 266	294-297	≥1.3	≥1%	FTS	1.21 (2x)
83Guelachv [47]	1191–1740	428 / 420	RT	13-1000	Nat	FTS	(2y)
00SiBeMaMa [76]	1480	1 / 1	RT	0.2-4.6	50%	L	2.7 (2z)
82PaCaFlGu [44]	2237–3114	1902/1892	295.8±1	1.3-8.9	≥10%	FTS	20.18-60.18 (2aa)
82ToGuBr [46]	2503–3099	575 / 574	296	21.4		FTS	0.25 (2bb)
92RiSmDeBe [59]	2576–2726	48 / 48	294-297	>1.3	>1%	FTS	1.21 (2cc)
83ToBr [48]	3289–4369	1002/1000	296	16.07	48%	FTS	(2dd)
73CaFlGuAm [41]	3702–3789	8/7				FTS	
07JeDaReTy [101]	4200–6599	3088/3074	293±2	2.7-22.7	Nat	FTS	
97Tothb [66]	4719–5781	1824/1816	RT	1.3-22.7	50%	FTS	0.25-73 (2ee)
07MiWaKaCa [103]	5964–7015	274 / 274		1.5-20		CRDS	(2ff)



01UIHuBeOn [81]	6003–7064	2424/2384	RT	15.16	15-44%	FTS	64-87	(2gg)
97Tothc [67]	6003–7626	2590/2571	296	1.8-13	3-49.1%	FTS	2.39-73	(2hh)
09LiNaKaCa [108]	6076–6623	498 / 498				CRDS		(2ii)
04MaRoMiNa [91]	6131–6748	1695/1693	RT	22.3	Nat	CRDS		(2jj)
89OhSa [55]	6384–6596	596 / 576	RT	12	50%	L		(2kk)
91SaTaIrNa [57]	6563–6565	3 / 3	RT	16.4	50%	L	1.0	(2ll)
05ToTe [92]	7428–8759	171 / 163	292 ± 3			FTS		(2mm)
04NaVoHu [90]	7500–8194	1415/1401	296	10	50%	FTS	87	
01HuHeZhWa [83]	8246-9063	899/873	RT	10	50%	FTS	105	(2nn)
05UIHuBeZh [95]	8713–9598	1000/996	296	18.4		FTS	123	
06NaLeCa [98]	9099–9645	1999/1900	RT	18	50%	ICLAS	≤23400	
08PeSoSi [106]	9387–9450	232 / 220		13-26				
05ToNaZoSh [93]	9758–10797	212 / 201	294.4	20.08	Nat	FTS	≤600	(2oo)
00WaHeHuZh [79]	9648–10201	654/649	RT	10	50%	FTS	123	(2pp)
03BeNaCa [87]	9628–10111	1416/1361	RT	26.3	50%	ICLAS	≤35000	
08ToTe [107]	9879–13901	133/87	RT	9-23	Nat	FTS	5-800	(2rr)
04NaHuHeCa [89]	10110–12215	1827/1787	RT	15	50%	FTS	105	(2ss)
97VoFaPIRi [65]	10631–10690	17 / 16				OPO		
08LiHo [105]	10693–10738	4/4	RT		Nat	CRDS		(2tt)
07MaToCa [102]	11809–12126	64 / 56	RT	26.7		ICLAS	4800-14000	(2uu)
05CaVaNa [96]	11645–12329	1615/1542	RT	18.4	50%	ICLAS	9000	
07VoNaCaCo [100]	11750–22708	3481/3378	291	6.7	48%	FTS	600	(2vv)
08NaVoMaTe [104]	12144–13159	1330/1300	RT	18.4	50%	ICLAS	≤24000	
99HuLiHeCh [70]	12525–12912	243/219				FTS		(2ww)
98LaPeSiZh [68]	12720–12912	121 / 120	300	17	50%	PAS		(2xx)
10NaBeLeCa [109]	13021–14113	2078/ 2019	RT	18.4	various	ICLAS	24000	
00NaBeCaSc [75]	13165–13500	676 / 572	RT	18.4	50%	ICLAS	≤15000	
99NaBeCa [71]	13579–14033	459 / 391	298	26.7	50%	ICLAS	≤15000	
00NaCa [73]	14986–15340	470 / 405	RT	17.3	50%	ICLAS	≤12000	
00CaBeNa [74]	16304–18349	400 / 348	RT			50% ICLAS		
01JeMeCaCo [80]	16344–22703	507 / 488	291.1	13.07	47%	FTS	603.32	(2yy)
00BeNaCa [78]	16542–17054	309 / 287	RT	17.3	50%	ICLAS	12000,19500	
90ByKaKoNa [56]	16746–17013	237 / 203	RT	20	50%	PAS		

<sup>a</sup> The tags listed are used to identify experimental data sources throughout this paper. The range given represents the range corresponding to wavenumber entries within the MARVEL input file of the particular water isotopologue and not the range covered by the relevant experiment. Uncertainties of the individual lines can be obtained from the Supplementary Ma-

terial. Trans. = transitions, with A = number of assigned transitions in the original paper, V = number of transitions validated in this study.  $T$  = temperature (K), given explicitly when available from the original publication, with RT = room temperature.  $p$  = pressure (hPa). Abun. = abundance (%) of the given isotopologue in the gas mixture, with Enr = enriched and Nat = natural abundance. Rec. = experimental technique used for the recording of the spectrum, with SMM = Stark-modulated microwave spectrometer, L = different lasers, CCC = collisionally cooled cell, MC = microwave cavities, TMS = tunable microwave spectrometer, PAS = photo-acoustic spectrometer, and OPO = optical parametric oscillator.

*2.8. Comments on the data sources for HD<sup>16</sup>O given in Table 2*

(2a) 62TrBe A spectrometer containing a Stark-modulated microwave coaxial cavity was used to measure nine hyperfine components of two rotational transitions, the  $3_{31} \rightarrow 3_{30}$  and  $5_{42} \rightarrow 5_{41}$  lines at 487 and 825 MHz, respectively. The hyperfine components are not included in the MARVEL analysis.

(2b) 55WeBeHe The apparatus of Ref. [21] was used for the measurements. Frequencies and absorption coefficients of six transitions between 486 and 3045 MHz are reported. Frequencies of five of these have been previously measured in Ref. [21].

(2c) 53BeWe The spectrometer employed Stark modulation and operated at frequencies lower than 600 MHz.

(2d) 54PoSt The number of original measured line frequencies is 5. In addition, the authors reproduce three frequencies from two other measurements: one transition from 49Mcafee [17] and two transitions from “Microwave Spectroscopy Laboratory, Research Laboratory of Electronics, Massachusetts Institute of Technology (unpublished work)”. Note also that one of the five frequencies of this source ( $8577.7 \pm 0.1$  MHz, the  $7_{44} \rightarrow 7_{43}$  transition) was previously reported in 53BuSt [23] but without an uncertainty.

(2e) 68VeBlDy The hyperfine structure of one rotational line, at 8577.812

MHz, was investigated with a beam maser spectrometer. The wrong rotational assignment,  $7_{43} \rightarrow 7_{45}$ , of the reported line was changed to the correct  $7_{44} \rightarrow 7_{43}$ . Measured and calculated values of the hyperfine coupling constants are reported but not used here.

(2f) 53Crawford Three observed lines are reported in this paper. The  $8\,577.80 \pm 0.08$  and  $8\,837.22 \pm 0.12$  MHz transitions correspond to HD<sup>16</sup>O transitions  $7_{44} \rightarrow 7_{43}$  and  $10_{56} \rightarrow 10_{55}$ , respectively; only the latter was included in the final transitions file. The third one at 8884.83 MHz is neither an HDO nor a D<sub>2</sub>O line.

(2g) 57Posener Fine structure of the  $2_{21} \rightarrow 2_{20}$  rotational line reported previously in Ref. [18] was examined by Stark modulation spectroscopy. Five components were observed but not employed here.

(2h) 64ThKrLo Nine hyperfine components of the  $2_{21} \rightarrow 2_{20}$  rotational transition have been reported at about 10 278.2455 MHz with a beam maser spectrometer.

(2i) 67BIVeDy The hyperfine structure of the  $2_{21} \rightarrow 2_{20}$  rotational line at 10 278.246 MHz was investigated with a beam maser spectrometer. Measured values of the hyperfine coupling constants as well as the electric quadrupole coupling tensor are reported but not employed here.

(2j) 49Strandbe Four rotational absorption transitions,  $2_{21} \rightarrow 2_{20}$ ,  $3_{22} \rightarrow 3_{21}$ ,  $4_{14} \rightarrow 3_{21}$ , and  $5_{33} \rightarrow 5_{32}$ , have been observed and identified using the Stark effect. The frequencies of the first three transitions are given in the paper. The last one was reported earlier in Ref. [16].

(2k) 46ToMe Frequency of one pure rotational transition,  $5_{33} \rightarrow 5_{32}$ , was reported. The frequency value,  $22\,309 \pm 5$  MHz, is rather far from the more precise values reported later for the same transition.

(2l) 53JeBiMa The frequencies of four rotational lines have been observed in microwave absorption. One of them,  $26\,880.47 \pm 0.1$  MHz, was observed previously [17]. Two lines at  $24\,884.85 \pm 0.1$  (A) and  $26\,880.47 \pm 0.1$  (B) MHz correspond to the HD<sup>16</sup>O transitions  $8_{45} \rightarrow 8_{44}$  and  $6_{24} \rightarrow 7_{17}$ , respectively.

The two unassigned lines at  $30\,182.57 \pm 0.1$  (C) and  $30\,778.62 \pm 0.1$  MHz (D) correspond to  $D_2^{16}O$  transitions  $9_{63} \rightarrow 8_{72}$  and  $9_{64} \rightarrow 8_{71}$ , respectively. Stark spectra of all four lines and Zeeman spectra of the A and B lines were recorded.

(2m) 93GoFeDeDu Three microwave transitions of  $HD^{16}O$ , reported previously by 84MeDeHe [49], were broadened by  $H_2$ ,  $N_2$ ,  $O_2$ , and He for  $100 \leq T \leq 600$  K. Measurements made at 10 mTorr of  $HD^{16}O$  and about 40 broadening gas pressures up to 1 Torr.

(2n) 87BaAlAlPo Three rotational transitions in the (0 1 0) vibrational state are reported. Only one of them,  $4_{14} \rightarrow 3_{21}$  at 230 427.34 MHz, is confirmed. The other transitions,  $6_{33} \rightarrow 7_{26}$  at 334 018.30 MHz and  $5_{24} \rightarrow 4_{13}$  at 345 795.94 MHz, have very large discrepancies as compared to the expected frequency values (about 335 138 and 346 620 MHz, respectively).

(2o) 84MeDeHe 32 newly reported rotational transitions complemented by transitions reported previously [40].

(2p) 76FlGi Among the 20  $HD^{16}O$  transition only 13 were validated. The 7 excluded transitions have differences between reported and expected frequency values larger than  $0.02\text{ cm}^{-1}$ .

(2q) 85Johns Far infrared pure rotational spectra of six water isotopologues were recorded with pressures less than 1 Torr (only for wavenumbers below  $30\text{ cm}^{-1}$  was the pressure raised to 2 Torr) to minimize pressure shifts and broadening. The list of  $HD^{16}O$  transitions includes 84 MW frequencies from 84MeDeHe [49] and 212 measured line positions which correspond to 225 transitions. The experimental uncertainty was raised to  $0.0002\text{ cm}^{-1}$  instead of  $0.00005\text{ cm}^{-1}$  stated originally.

(2r) 78KaKaKy Three spectra were recorded but only for one spectrum ( $150\text{--}320\text{ cm}^{-1}$ ,  $p = 1.8$  Torr) was the  $HD^{16}O$  abundance greater than natural. The experimental accuracy has been estimated in the original source to be better than  $0.001\text{ cm}^{-1}$  under favorable conditions. After recalibration, 65% (39 lines) of the 65 observed transitions agree with the MARVEL

transitions within the declared experimental accuracy of  $1.0 \times 10^{-3} \text{ cm}^{-1}$ . All original experimental uncertainties equal to  $5 \times 10^{-4}$  were changed to  $1.5 \times 10^{-3}$ .

(2s) 95PaHo Pure rotational spectra of  $^{16}\text{O}$  water isotopic species were measured with a high resolution FTS in order to obtain better calibration line positions in the  $110 - 500 \text{ cm}^{-1}$  spectral region.

(2t) 03JaTeBeZo Fourier transform emission spectra analyzed together with line positions obtained in 01PaBeZoSh [82]. Supplementary data for this article contains a list of 32 399 records which correspond to 31 645 unique line positions. 10 657 rotation-vibration assignments are given in this line list. 745 of them are labeled with “ $\text{D}_2\text{O}$ ” or “ $\text{H}_2\text{O}$ ”, probably they are HDO lines blended by “ $\text{D}_2\text{O}$ ” and “ $\text{H}_2\text{O}$ ”, respectively. About 1 600 transitions which represent the second components of quasi-degenerate transitions, omitted in the original source, have been added to the list, as well as 65 newly assigned transitions. After validation of the spectrum, a list of 11 182 transitions has been included into the final transitions file. Of the 11 182 validated transitions, only 67.9% (7 592 lines) agree with the MARVEL transitions within the declared experimental uncertainty of  $1.0 \times 10^{-3} \text{ cm}^{-1}$ , suggesting that the true experimental accuracy of the line positions can be lower than declared. Furthermore, all original experimental uncertainties of  $5.0 \times 10^{-3}$  have been lowered to  $2.0 \times 10^{-3} \text{ cm}^{-1}$ .

(2u) 99Toth High-resolution Fourier-transform spectra were analyzed to obtain line positions and strengths of over 6 000 transitions of HDO and  $\text{D}_2\text{O}$ . For  $\text{HD}^{16}\text{O}$ , line positions and strengths of pure lines and of rovibrational bands  $(0\ 1\ 0)-(0\ 0\ 0)$ ,  $(0\ 2\ 0)-(0\ 1\ 0)$ , and  $(1\ 0\ 0)-(0\ 1\ 0)$  are reported. Of the 749 observed transitions reported for  $\text{HD}^{16}\text{O}$ , 95.5% (715 lines) agree with the MARVEL transitions within the declared experimental uncertainty of  $1 \times 10^{-4} \text{ cm}^{-1}$ . There is an asymmetry in the obs - MARVEL deviations with respect to the zero line: there are 276 positive and 473 negative deviations. This asymmetry concerns mainly the  $1\ 000 - 1\ 360 \text{ cm}^{-1}$  region. A

calibration factor could not be determined.

(2v) 93Tothb Fourier transform spectra in the  $\nu_2$  band region have been recorded for the determination of line positions and line intensities of HD<sup>16</sup>O, HD<sup>17</sup>O, and HD<sup>18</sup>O. Of the 1 267 lines assigned to HD<sup>16</sup>O, 94.8% (1 201 lines) agree with MARVEL transitions within the declared experimental uncertainty of  $1.2 \times 10^{-4} \text{ cm}^{-1}$ . An asymmetry is detected in the obs – MARVEL deviations with respect to zero line: there are 412 positive and 855 negative deviations. A calibration factor could not be determined as it does not seem to be constant for the whole spectral region considered.

(2w) 86FlCaMaGu Four spectra recorded between 1 040 and 1 900  $\text{cm}^{-1}$  with different pressures and mixtures of D<sub>2</sub>O and H<sub>2</sub>O in order to obtain  $\nu_2$  line positions of HD<sup>16</sup>O. After recalibration of the experimental spectrum, 99.3% (1 244 lines) of the 1 253 observed transitions agree with the MARVEL transitions within  $2.0 \times 10^{-4} \text{ cm}^{-1}$ , compared to the stated experimental uncertainty of  $4.0 \times 10^{-4} \text{ cm}^{-1}$ . The experimental uncertainties of the lines between 1 210 and 1 890  $\text{cm}^{-1}$  were decreased from  $4 \times 10^{-4}$  to  $2 \times 10^{-4} \text{ cm}^{-1}$ . As the applied calibration factor was not constant throughout the spectrum, there is a remaining asymmetry in the obs – MARVEL deviations with respect to the zero line: there are 764 positive and 489 negative deviations. In the region below 1 200  $\text{cm}^{-1}$  most part of the observed – MARVEL deviations are negative, with a root-mean-square agreement of  $3.6 \times 10^{-4} \text{ cm}^{-1}$ . Uncertainties that belong to the 1 210–1 890  $\text{cm}^{-1}$  interval were changed from  $4.0 \times 10^{-4} \text{ cm}^{-1}$  to  $2.0 \times 10^{-4} \text{ cm}^{-1}$ .

(2x) 91RiSmDeBe Measured Lorentz-broadening and pressure-shift coefficients by air, N<sub>2</sub>, and O<sub>2</sub> are presented for the 266  $\nu_2$  HD<sup>16</sup>O lines. Twelve spectra were recorded at 0.0053  $\text{cm}^{-1}$  resolution. Spectra with buffer gas were recorded at total sample pressures of about 200, 300, and 400 Torr. Spectra have been calibrated using 93Tothb [61] data, and then the experimental accuracy has been adopted in this study as  $1.2 \times 10^{-4} \text{ cm}^{-1}$ , similar to that of 93Tothb. All 265 observed lines agree with MARVEL transitions

within  $7.4 \times 10^{-5} \text{ cm}^{-1}$ . An asymmetry in the obs – MARVEL deviations with respect to the zero line: there are 182 positive and 83 negative deviations. A calibration factor could not be determined with the technique employed in this study as it is not constant for the whole spectral region considered. Original experimental uncertainties of  $5.0 \times 10^{-4} \text{ cm}^{-1}$  have been adjusted to  $1.2 \times 10^{-4} \text{ cm}^{-1}$ .

(2y) 83Guelachv Fourier transform spectra with absorption paths of 16, 24.17, 32.17, and 44.17 m. Due to a calibration problem the original reported transitions had to be scaled by a scale factor of 0.999 999 77 [130]. The MARVEL input file includes only the recalibrated transitions, with the original uncertainties. After recalibration of the spectrum, 95.1% (407 lines) of the 428 observed transitions agree with the MARVEL transitions with an rms deviation of  $1.5 \times 10^{-4} \text{ cm}^{-1}$ , compared to an experimental uncertainty of  $1.2 \times 10^{-4} \text{ cm}^{-1}$ , estimated as an average value of individual experimental uncertainties attached to every line. As the applied calibration factor was not constant throughout the spectrum, there is a remaining asymmetry in the obs – MARVEL deviations with respect to the zero line: there are 151 positive and 277 negative deviations, with the most serious discrepancies in the 1 400–1 600  $\text{cm}^{-1}$  region.

(2z) 00SiBeMaMa A CO laser was used for observation of the saturated absorption dip of the nearby  $\nu_2$   $5_{24} \rightarrow 5_{15}$  line of HD<sup>16</sup>O. The absolute frequency of this line was obtained as  $1\,480.094\,038\,033 \pm 67 \times 10^{-9} \text{ cm}^{-1}$  ( $44\,372\,102.973 \pm 0.002 \text{ MHz}$ ).

(2aa) 82PaCaFlGu The  $2\nu_2$  and  $\nu_1$  bands of HD<sup>16</sup>O were analyzed using Fourier transform spectra of pure D<sub>2</sub>O and of 90% H<sub>2</sub>O – 10% D<sub>2</sub>O mixtures.

(2bb) 82ToGuBr HDO gas samples obtained from different mixtures of H<sub>2</sub>O and D<sub>2</sub>O with concentrations of HD<sup>16</sup>O estimated as 50, 48, and 28%. For a number of lines the experimental uncertainties were adjusted from  $4 \times 10^{-3}$  to  $4 \times 10^{-4} \text{ cm}^{-1}$ .

(2cc) 92RiSmDeBe Absorption spectra recorded for measurements of Lorentz-broadening coefficients and pressure-induced line-shift coefficients of 48 lines in the  $\nu_1$  band of HDO were recorded using the FTS of the Kitt Peak National Observatory. Gas samples were prepared by mixing distilled H<sub>2</sub>O with 99.96 atom% D<sub>2</sub>O. Low pressure ( $< 1.0$  Torr) samples were mixed with samples diluted with high-purity N<sub>2</sub>, O<sub>2</sub>, or ultra zero air at total sample pressures of about 200, 300, and 400 Torr.

(2dd) 83ToBr FTS spectra recorded by 82ToGuBr [46] [see comment (2bb)] were studied for line positions and strengths in the 3 289–4 369 cm<sup>-1</sup> region.

(2ee) 97Tothb The FTS of the Kitt Peak National Observatory was used for measuring HD<sup>16</sup>O spectra with absorption path lengths of 0.25, 2.39, 25, and 73 m. Of the 1 824 observed transitions, 66.5% (1 213 lines) agree with the MARVEL transitions with an rms deviation equal to the declared experimental uncertainty of  $1.0 \times 10^{-4}$  cm<sup>-1</sup>. The measurement accuracy seems to be slightly overestimated.

(2ff) 07MiWaKaCa Absorption spectra of natural abundance water vapor near the 5911.0–5922.5, 5926.0–5941.8, 5957.0–6121.6, and 6475.0–7015.6 cm<sup>-1</sup> spectral regions. Pressure values on the order of 10–20 hPa were chosen in the 1.66  $\mu$ m region and of about 1.5–2.0 hPa in the 1.455  $\mu$ m region. Most of the original experimental uncertainties were increased, from  $8.0 \times 10^{-5}$  cm<sup>-1</sup> to  $1.0 \times 10^{-3}$  cm<sup>-1</sup>.

(2gg) 01UIHuBeOn The Bruker IFS 120 HR FTS of the University of Science and Technology (Hefei, China) was used to record two spectra: the first spectrum had a total pressure of 1 516 Pa, 44% HDO and a path length of 87 m; the second spectrum had a total pressure of 1 500 Pa, 15% HDO and a path length of 69 m. Transition data supplied by S-M Hu. Of the 2 424 observed transitions, 92.9% (2 251 lines) agree with the MARVEL transitions with an rms deviation equal to the declared measurement accuracy of  $2.0 \times 10^{-3}$  cm<sup>-1</sup>. There is an asymmetry in the observed–MARVEL deviations with respect to the zero line: there are 735 positive and 1689 negative deviations.



A calibration factor for the whole spectral region covered could not be determined. For 60% of the reported transitions  $\text{obs} - \text{MARVEL}$  deviations are less than  $0.001 \text{ cm}^{-1}$ , confirming that the measurement accuracy is slightly better than declared.

(2hh) 97Tothc Of the 2 590 observed transitions, 61.8% (1 600 lines) agree with the MARVEL transitions with an rms deviation equal to the declared experimental uncertainty of  $1.0 \times 10^{-4} \text{ cm}^{-1}$ . The stated measurement accuracy seems to be slightly overestimated.

(2ii) 09LiNaKaCa Absorption spectrum of  $^{18}\text{O}$  enriched water was recorded using CW-CRDS and a series of fibered DFB lasers. The CRDS sensitivity allowed the detection of lines with intensity as low as  $10^{-28} \text{ cm/molecule}$ . Uncertainties of  $2.0 \times 10^{-3} \text{ cm}^{-1}$  were adjusted to  $1.0 \times 10^{-3} \text{ cm}^{-1}$ .

(2jj) 04MaRoMiNa A pressure of 17 Torr was used in the whole spectral region; additional recordings at 1 Torr pressure were performed above  $6\,510 \text{ cm}^{-1}$  due to the presence of strong lines. As the  $\text{HD}^{16}\text{O}$  absorption lines have been observed in natural abundance, they mostly represent weak lines. Thus, the declared measurement accuracy, better than  $0.001 \text{ cm}^{-1}$  on average, for these lines was decreased to  $0.002 \text{ cm}^{-1}$ . Of the 1 695 observed transitions, 88.7% (1 503 lines) agree with MARVEL transitions with an rms deviation equal to the adopted measurement accuracy of  $2 \times 10^{-3} \text{ cm}^{-1}$ . All transitions with measurement accuracy of  $3 \times 10^{-3} \text{ cm}^{-1}$  have been changed to  $2 \times 10^{-3} \text{ cm}^{-1}$ .

(2kk) 89OhSa Absorption spectra of HDO and  $\text{D}_2\text{O}$  recorded using single-mode distributed feedback semiconductor lasers. After recalibration of the experimental spectrum, 96.3% (574 lines) of the 596 observed transitions agree with the MARVEL transitions within an rms deviation of  $4.4 \times 10^{-3} \text{ cm}^{-1}$ , as compared to the measurement accuracy of  $5.0 \times 10^{-3} \text{ cm}^{-1}$ .

(2ll) 91SaTaIrNa A near-infrared semiconductor-laser spectrometer was used for measurements of absolute frequencies of 30 vibration-rotation transitions of several molecular species. Three HDO frequencies were measured with an

accuracy of about 30 MHz and can be used as a calibration standard in the near infrared region. Vibrational labeling given in the paper was changed by comparison with other assignments: the band studied is not (2 0 1)–(0 0 0) but (0 2 1)–(0 0 0).

(2mm) 05ToTe Experimental spectra recorded by Schermaul *et al.* [136] were analysed here.

(2nn) 01HuHeZhWa Transitions list supplied by S.-M. Hu. Of the 899 observed transitions, 85.6% (770 lines) agree with the MARVEL transitions with an rms deviation equal to the declared measurement accuracy of 0.001  $\text{cm}^{-1}$ .

(2oo) 05ToNaZoSh Re-analysis of the absorption spectra of natural abundance water vapor in the 9 250–26 000  $\text{cm}^{-1}$  region reported previously by Coheur *et al.* [84] and by Mérienne *et al.* [88]. The original spectra were recorded by an FTS Bruker IFS 120M coupled with two White multiple-reflection cells of 5 and 50 m base path. The total absorption path was up to 600 m. For weak lines assigned to HD<sup>16</sup>O absorption, an experimental accuracy of 0.003  $\text{cm}^{-1}$  has been adopted in this study, though in the original source the measurement accuracy was declared to be higher (on average). Of the 212 transitions assigned to HD<sup>16</sup>O absorption, only 58% (212 lines) agree with the MARVEL transitions with an rms deviation equal to the adopted experimental accuracy of 0.003  $\text{cm}^{-1}$ , suggesting that the measurement accuracy was indeed slightly overestimated.

(2pp) 00WaHeHuZh Spectra recorded with the Bruker IFS 120HR FTS of the University of Science and Technology (Hefei, China). Transitions list supplied by S.-M. Hu. Of the 626 observed transitions, 98.2% (642 lines) agree with the MARVEL transitions with an rms deviation equal to the declared measurement accuracy of 0.003  $\text{cm}^{-1}$ . There is an asymmetry in the observed–MARVEL deviations with respect to the zero line: there are 248 positive and 406 negative deviations. A calibration factor for the whole spectral region could not be determined. For 91% of the reported transitions

observed–MARVEL deviations are within  $0.0016 \text{ cm}^{-1}$ , suggesting that the measurement accuracy is better than declared.

(2rr) 08ToTe A multi-pathlength refit of the data of Schermaul *et al* [136, 137].

(2ss) 04NaHuHeCa A 1 : 1 mixture of  $\text{H}_2\text{O}$  and  $\text{D}_2\text{O}$  was used to make HDO. An additional spectrum of water vapor with much higher concentration of  $\text{D}_2\text{O}$  was recorded. The observed spectrum consists of two parts: 10 110–11 451 and 11 751–12 215  $\text{cm}^{-1}$ , containing 1 489 and 338 lines, respectively. Of the 1 489 observed transitions in the first region, 85.8% (1 277 lines) agree with the MARVEL transitions with an rms deviation equal to the declared measurement accuracy of  $0.002 \text{ cm}^{-1}$ . In the 11 751–12 215  $\text{cm}^{-1}$  region there is an asymmetry in the observed–MARVEL deviations with respect to the zero line: there are 252 positive and 84 negative deviations. A calibration factor for this spectral region could not be determined. For 75.4% of the transitions observed–MARVEL deviations are within  $0.005 \text{ cm}^{-1}$ , suggesting that the experimental accuracy is lower in this region.

(2tt) 08LiHo CRDS was used for precision measurements of the water line intensities in the 10 603–10 852  $\text{cm}^{-1}$  region. Four lines of the  $3\nu_3$  band of HDO were measured.

(2uu) 07MaToCa ICLAS absorption spectra of  $^{18}\text{O}$  enriched water were recorded; HDO lines were observed due to high sensitivity of the spectrometer used.

(2vv) 07VoNaCaCo Assignment of previously reported FTS spectra of Bach *et al.* [94].

(2ww) 99HuLiHeCh After applying a calibration factor of 1.000 000 70 (Table 1), 86% (209 lines) of the 243 observed transitions agree with the MARVEL transitions with an rms deviation equal to the declared measurement accuracy of  $0.005 \text{ cm}^{-1}$ .

(2xx) 98LaPeSiZh Absorption spectrum of HDO was investigated using a photo-acoustic spectrometer based on a Ti-Sapphire laser.

(2yy) 01JeMeCaCo Some of the lines published in this paper were assigned by 07VoNaCaCo [100]. Of the 507 observed transitions, 84.6% (429 lines) agree with MARVEL transitions with an rms deviation equal to the declared measurement accuracy of  $0.002 \text{ cm}^{-1}$ .

Table 3: Data sources and their characteristics for HD<sup>17</sup>O. See section 2.9 for comments.<sup>a</sup>

Tag	Range (cm <sup>-1</sup> )	Trans. A / V	Physical conditions				Comments
			<i>T</i> (K)	<i>p</i> (hPa)	Abun.	Rec.	
78Lovas [110]	0.29–8.06	59/54	RT				(3a)
57StTo [111]	0.346	2 /2	RT				(3b)
93Tothb [61]	1118–1674	424/422	RT			FTS	(3c)

<sup>a</sup> See footnote a to Table 2.

### 2.9. Comments on the data sources for HD<sup>17</sup>O given in Table 3

(3a) 78Lovas 59 microwave transitions reproduced from G. Steenbeckeliers, private communication, July 1971. This data source contains three orphans: 78Lovas.2 at 0.305 972, 78Lovas.11 at 0.781 876, and 78Lovas.45 at 4.614 810 cm<sup>-1</sup>. 78Lovas.27 is also part of an FSN along with 93Tothb.72.

(3b) 57StTo The quadrupole hyperfine structure of the 2<sub>20</sub> → 2<sub>21</sub> rotational transition was measured. The unsplit line position is given as 10374.56 MHz, while the strongest component of the multiplet is estimated as 10374.31 MHz.

(3c) 93Tothb 93Tothb.72 is part of an FSN along with 78Lovas.27.

Table 4: Data sources and their characteristics for HD<sup>18</sup>O. See section 2.10 for comments.<sup>a</sup>

Tag	Range (cm <sup>-1</sup> )	Trans. A / V	Physical conditions					Comments
			<i>T</i> (K)	<i>p</i> (hPa)	Abun.	Rec.	<i>L</i> (m)	
78Lovas [110]	0.20–8.06	27/ 27	RT					(4a)
06ZeBaKuRi [112]	16.73	1 / 1	RT					(4b)
76FIGi [42]	16.4–39.9	17/ 17	RT					(4c)
85Johns [51]	24.9–219.8	191/ 191	RT					(4d)
06LiDuSoWa [138]	1090–7609	4931/4896		13.07	0.8%	FTS	105	(4k)
93Tothb [61]	1125–1804	704/ 702	RT	1.60	0.86%	FTS	2.4	(4e)
05Toth [113]	2517–3018	354/ 352	RT	2.00	10.5%	FTS	2.4	(4f)
09MiTaPuSt [114]	4205–10717	1012/1003	RT	3–19	Nat	FTS	600–1800	(4h)
10MiTaDaJe [115]	5600–8744	299/298	296	3–13	Nat	FTS	600	(4l)
09LiNaKaCa [108]	6000–6722	1101/1097	299	0.13–13	0.028	CRDS		(4i)
07MaToCa [102]	11638–12105	92/ 92	RT	26.7		ICLAS	4800–14000	(4j)

<sup>a</sup> See footnote a to Table 2.

### 2.10. Comments on the data sources for HD<sup>18</sup>O given in Table 4

(4a) 78Lovas 27 microwave transitions reproduced from G. Steenbeckeliers, private communication, July 1971 (seemingly the first observation of the HD<sup>18</sup>O spectrum). The correct assignment of the pure rotational transition reported at 5 902.38 MHz is [4 3 1] ← [4 3 2] instead of [4 4 1] ← [4 3 2].

(4b) 06ZeBaKuRi The line of the [1 1 0] ← [1 0 1] transition has been detected at a frequency of 501 567.39 MHz in the atmosphere using the sub-millimeter radiometer (SMR) instrument aboard the Odin satellite.

(4c) 76FIGi Source commented on in detail in 09MiTaPuSt [114], see its Sections 2 and 6 and Table 4. One of the first studies of water vapor spectra using the Fourier transform technique.

(4d) 85Johns Source commented on in detail in 09MiTaPuSt[114], see its

Sections 2 and 6 and Table 4. Note that 24 MW wavenumbers reproduced by Johns from 78Lovas [110] were incorrectly cited in 83MeLuHe [139].

(4e) 93Tothb See section 2.2 for comments about re-calibration.

(4f) 05Toth Transitions corresponding to the  $2\nu_2$  and  $\nu_1$  bands. For more detailed comments see Sections 2 and 6 of 09MiTaPuSt.

(4h) 09MiTaPuSt Assignment of 38 unassigned lines reported by 07JeDaReTy [101] in the 4 205–6 588  $\text{cm}^{-1}$  range (see Section 3.1 and Table 1 of 09MiTaPuSt) and 970 transitions from 6 129 to 10 717  $\text{cm}^{-1}$ . The transition  $(0\ 2\ 1)[6\ 2\ 4] \leftarrow (0\ 0\ 0)[5\ 2\ 3]$  transition reported at 6 538.243 39  $\text{cm}^{-1}$  is re-assigned to 6 538.261 39  $\text{cm}^{-1}$ .

(4i) 09LiNaKaCa This source contains two orphans:

6 139.12(0.001)  $(1\ 0\ 1)[14\ 1\ 13] \leftarrow (000)[15\ 2\ 14]$ , 09LiNaKaCa.43 and

6 139.6(0.001)  $(1\ 0\ 1)[14\ 2\ 13] \leftarrow (0\ 0\ 0)[15\ 1\ 14]$ , 09LiNaKaCa.44.

The following transitions have been deleted by the MARVEL analysis:

6 157.76  $(1\ 0\ 1)[13\ 2\ 12] \leftarrow (0\ 0\ 0)[14\ 1\ 13]$ , 09LiNaKaCa.54 and

6 538.88  $(1\ 0\ 1)[13\ 2\ 12] \leftarrow (0\ 0\ 0)[12\ 1\ 11]$ , 09LiNaKaCa.945. One transition, 09LiNaKaCa.594 is part of an FSN.

(4j) 07MaToCa Transitions belonging to the  $\nu_2 + 3\nu_3$  band are commented in 09MiTaPuSt [114], see Sections 2 and 6 and Figure 4.

(4k) 06LiDuSoWa 17 transitions are part of an FSN, 7 are ORPs and 13 have been deleted by the MARVEL analysis.

(4l) 10MiTaDaJe The transition deleted by MARVEL is: 7 015.213 870  $(0\ 0\ 2)[12\ 1\ 11] \leftarrow (0\ 0\ 0)[13\ 1\ 12]$ , 10MiTaDaJe.186.

### 3. MARVEL energy levels

Tables 5–7 contain MARVEL vibrational band origins (VBO) for HD<sup>16</sup>O, HD<sup>17</sup>O, and HD<sup>18</sup>O, respectively. The same tables also contain information about the number of rovibrational energy levels validated within this work for each VBO. Observed, MARVEL predicted, and variationally determined spectra of HD<sup>16</sup>O (VTT [118]), HD<sup>17</sup>O [116, 140], and HD<sup>18</sup>O [116, 140] are reported in Figs. 4-6, respectively.

#### 3.1. HD<sup>16</sup>O

For HD<sup>16</sup>O, 53 291 transitions of the 54 740 initial transitions have been validated and used in the final MARVEL analysis. From these transitions we derive a final set of 8 818 energy levels which belong to 54 vibrational states and have  $J$  up to 30 and  $K_a$  up to 21. 66 transitions from 03JaTeBeZo [86] form an FSN and 13 transitions from the same source are orphans; thus, they could not be used for energy level determination.

A comparison of the experimental (MARVEL) and variational (VTT [118]) energy level values is given in Fig. 2. The root-mean-square deviation for this comparison is 0.058 cm<sup>-1</sup>, with the maximum obs – calc residuals being –0.51 and –1.27 cm<sup>-1</sup> for the [4 0 4] and [3 0 3] levels of the (0 6 3) vibrational state, respectively. The (0 6 3) [3 0 3] energy level at 18 293.160 7 cm<sup>-1</sup> was derived from two weak lines in 07VoNaCaCo [100] and should most likely be considered incorrect. Note that these two levels were excluded when the above quoted rms deviation was determined.

A comparison can be made between the original set of observed transitions and those calculated from the experimental energy levels determined by MARVEL; this is presented in Fig. 3. The differences do not exceed 0.038 cm<sup>-1</sup>; 52.6% and 84% of all transitions are reproduced within 0.001 and 0.005 cm<sup>-1</sup>, respectively. Only 5.6% of the observed transitions differ from the MARVEL calculated values by more than 0.01 cm<sup>-1</sup>.



MARVEL may increase, via robust reweighting, an experimental uncertainty of a transition when it is not consistent with that derived from the MARVEL energy levels. For transitions with low  $J$  and  $K_a$  values it is rather easy to evaluate a feasible experimental uncertainty if enough CDs are available. However, for increased values of  $J$  and  $K_a$  the CD relations become less accurate and instructive, a number of experimental lines represent unresolved multiplets, and it becomes more and more difficult to judge properly the actual experimental accuracy of the transitions. This in turn limits the accuracy of the MARVEL energy levels derived. This situation could be improved only by including additional accurate experimental information in the MARVEL input file.

The set of MARVEL energy levels derived from processing the validated observed transitions has been used to predict a large number of rovibrational transitions with positions at the level of experimental accuracy. These line positions were augmented with variational intensities. The total number of predicted transitions with intensities larger than  $1.0 \times 10^{-29}$  cm molecule<sup>-1</sup> is 188 565 for  $T = 296$  K. Observed, observed + predicted, and variational transitions are shown in Fig. 4 for HD<sup>16</sup>O. The number of residual unobserved variational lines was obtained as the difference between the variationally calculated and observed + predicted transitions; their number is 185 751 with intensities greater than  $1 \times 10^{-29}$  cm molecule<sup>-1</sup>. This figure is especially important for future experimental studies of the high-resolution spectra of HD<sup>16</sup>O. The strongest unobserved lines fall in the 3 600–9 600 and 9 600–12 400 cm<sup>-1</sup> regions with intensities  $3 \times 10^{-23}$  and  $4 \times 10^{-25}$  cm molecule<sup>-1</sup>, respectively. Interestingly, in the spectral region below 2 400 cm<sup>-1</sup> intensities of the residual lines do not exceed  $1 \times 10^{-27}$  cm molecule<sup>-1</sup>, and in the 2 400–3 600 cm<sup>-1</sup> region their intensities are not larger than  $1 \times 10^{-26}$  cm molecule<sup>-1</sup>. Such a complete coverage of the 0–2 400 cm<sup>-1</sup> region by the obs + predicted transitions is due to energy levels derived from the analysis of hot spectra in 03JaTeBeZo [86]. Similarly to the ‘cold’ spectrum, a set

of 74 122 ‘hot’ transitions at  $T = 1\,770$  K in the  $380\text{--}3\,932$   $\text{cm}^{-1}$  region have been predicted at a level of experimental accuracy in positions using MARVEL energy levels and variational intensities larger than  $1.0 \times 10^{-24}$   $\text{cm molecule}^{-1}$ .

We considered three sources of hot  $\text{HD}^{16}\text{O}$  transitions, Refs. [64, 82, 86]. The list given in 03JaTeBeZo [86] includes the transitions given by 01PaBeZoSh [82], so it was not necessary to analyse the second source independently. Since hot water spectra have been measured for the most abundant isotopologue,  $\text{HD}^{16}\text{O}$ , the largest  $J$  is high,  $J_{\text{max}} = 30$ . Then, one can observe that, due to the large number of measured transitions, the list of VBOs is fairly complete for  $\text{HD}^{16}\text{O}$  and there are very few VBOs below  $15\,000$   $\text{cm}^{-1}$  for which rotational-vibrational levels have not been determined. In fact, the only lower-lying VBOs which are missing are (1 4 0), (2 4 0), (0 9 0), (2 3 1), (1 3 2), and (2 1 2). At higher wavenumbers a high proportion of VBOs still appear (Table 5). The highest states for which VBOs and rovibrational levels have been measured contain 7 quanta of stretch: 51 and 40 rotational levels have been determined for the (0 0 7) and (1 0 6) VBOs, respectively. We note that double resonance measurements [97] give data on VBOs higher than these; however, this work presents no measured transition wavenumbers and we were therefore unable to use it in this study.

Table 5: MARVEL vibrational band origins (VBO) for HD<sup>16</sup>O, with normal-mode ( $v_1v_2v_3$ ) assignments, MARVEL uncertainties, and the number of validated rotational-vibrational levels (RL) associated with the vibrational levels in the present database.<sup>a</sup>

$v_1v_2v_3$	VBO/cm <sup>-1</sup>	Unc. <sup>a</sup>	RL
0 0 0	0.000000 <sup>b</sup>	0	685
0 1 0	1403.483724	15	609
1 0 0	2723.679737	49	478
0 2 0	2782.011177	36	457
0 0 1	3707.466740	177	380
1 1 0	4099.955912	250	347
0 3 0	4145.473186	292	312
0 1 1	5089.539837	41	253
2 0 0	5363.824480	50	209
0 4 0	5420.041442	1016	76
1 2 0	5506.186812	250	114
1 0 1	6415.460697	42	164
0 2 1	6451.899800	75	151
0 5 0	6690.413212	999	45
2 1 0	6746.908202	36	113
1 3 0	[6849.00]		11
0 0 2	7250.517890	257	205
0 3 1	7754.605467	500	110
1 1 1	7808.758612	500	103
0 6 0	7914.317012	500	11
3 0 0	7918.171912	500	164
2 2 0	[8090.15]		16
1 4 0	[8173.13]		

0 1 2	8611.100944	365	151
0 4 1	[9032.23]		111
2 0 1	9047.068467	500	146
0 7 0	[9086.21]		54
1 2 1	9155.817812	2499	108
3 1 0	9293.001312	536	180
1 5 0	9381.786458	497	64
2 3 0	9487.915506	491	86
0 2 2	9934.772127	873	150
1 0 2	9967.016071	1019	152
0 8 0	[10118.42]		6
0 5 1	[10318.47]		12
4 0 0	10378.945940	1224	89
2 1 1	10403.147412	999	19
1 3 1	[10480.59]		2
1 6 0	[10602.72]		1
0 0 3	10631.683126	619	172
3 2 0	[10652.58]		9
<i>2 4 0</i>	<i>[10789.66]</i>		
<i>0 9 0</i>	<i>[11110.69]</i>		
0 3 2	11242.923202	2499	36
1 1 2	11315.433312	999	54
0 6 1	[11533.36]		10
3 0 1	[11582.72]		86
2 2 1	11701.775912	2499	51
4 1 0	11754.581840	1783	81
1 7 0	11773.312167	2499	7
1 4 1	[11804.55]		11
3 3 0	[11958.60]		27
0 1 3	11969.753013	12	181

2 5 0	[12073.36]		1
0 10 0	[12164.86]		1
0 4 2	[12516.42]		26
2 0 2	12568.190090	1809	77
1 2 2	12644.652812	2499	57
0 7 1	[12694.33]		3
5 0 0	12767.141475	309	119
1 8 0	[12852.35]		3
3 1 1	12919.938712	2499	78
1 5 1	[12986.79]		9
<i>2 3 1</i>	<i>[13087.15]</i>		
0 2 3	13278.350811	46	141
1 0 3	13331.606153	216	126
0 8 1	[13716.47]		1
0 5 2	[13797.38]		5
0 0 4	13853.627343	155	185
<i>2 1 2</i>	<i>[13889.58]</i>		
<i>1 3 2</i>	<i>[13960.04]</i>		
5 1 0	[14147.42]		31
0 3 3	[14563.09]		8
0 12 0	[14564.83]		8
1 1 3	14660.721912	315	42
0 6 2	[14997.32]		1
6 0 0	15065.712212	4997	37
3 0 2	[15097.31]		2
0 1 4	15166.104512	17	130
1 4 2	15170.951012	33	30
4 1 1	[15349.50]		2
0 4 3	[15822.14]		1
<i>2 0 3</i>	<i>[15924.17]</i>		

1 11 0	[16047.64]		3
1 5 2	[16449.57]		1
0 2 4	16456.190313	60	75
5 0 1	[16481.72]		3
1 0 4	[16539.04]		42
0 0 5	16920.020712	15	135
1 1 4	[17844.35]		1
0 6 3	[18202.57]		2
0 1 5	18208.446512	48	85
<i>2 7 1</i>	<i>[18919.13]</i>		
0 2 5	[19472.44]		25
1 0 5	[19584.53]		21
3 1 3	[19742.09]		1
0 0 6	19836.882808	45	87
<i>7 0 1</i>	<i>[20952.89]</i>		
1 0 6	22454.468792	500	40
0 0 7	22625.523004	2161	51
<i>1 1 6</i>	<i>[23700.84]</i>		

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<sup>a</sup> The uncertainties (Unc.) are given in units of  $10^{-6} \text{ cm}^{-1}$ . For VBOs not determined by the available experimental data, approximate variationally computed VBOs, based on an exact kinetic energy operator and the PES of Yurchenko *et al.* [119], are given in brackets. These values should only be used for guidance about the VBOs, although their accuracy is expected to be better than  $0.1 \text{ cm}^{-1}$ . No uncertainties are given for these VBOs. For completeness, VBOs for which no rovibrational states have been observed are also given. These levels, for which no RL values are given, are printed in italics and correspond to variationally computed values, as well. The VBOs

are ordered according to energy.

<sup>b</sup> The value of the vibrational ground state was fixed to zero with zero uncertainty.

### 3.2. $HD^{17}O$ and $HD^{18}O$

Of the 485 transitions in the MARVEL database of  $HD^{17}O$ , 478 were validated and these led to 162 MARVEL energy levels with a maximum  $J$  value of 11. There is only one VBO determined for  $HD^{17}O$ , (0 1 0) (Table 6 and Figure 5). As is clear from Fig. 5, the energy levels determined for this band allow one to make excellent predictions of a large number of pure rotational levels on the ground vibrational state (0 0 0). Future experimental investigations should validate these MARVEL predictions.

As to  $HD^{18}O$ , of the 8 729 transitions in the MARVEL database, 8 634 were validated and these led to the determination of 1 860 MARVEL energy levels with a maximum  $J$  value of 18. Several recent studies have addressed the assignment of the high-resolution rovibrational spectrum of  $HD^{18}O$ , including 06LiDuSoWa [138], 09MiTaPuSt [114], 09LiNaKaCa [108], and 10MiTaDaJe [115]. While there have been a relatively small number of assigned transitions for  $HD^{18}O$  before, 06LiDuSoWa changed this situation drastically. Furthermore, the experimental database of  $HD^{18}O$  was treated by the code RITZ in 09MiTaPuSt [114] and 10MiTaDaJe [115]. The RITZ code is similar to MARVEL and results in energy levels and their uncertainties. Table 7 compares the MARVEL and the RITZ VBO and uncertainty values. The recommended values remain those obtained using the present database and the MARVEL protocol.

### 3.3. Status of highly accurate transitions

The agreement between the MARVEL predicted and the experimental pure rotational transitions improved slightly by the recalibration of the experimental transitions. Nevertheless, the MARVEL uncertainties of the pure rotational levels is still uniformly larger than is usual for lines coming from microwave determinations. To show that this is due to the uncertainties of the upper states we performed a MARVEL analysis of the pure rotational states. MARVEL can reproduce the microwave uncertainties very nicely, down to the level of the experimental uncertainties, as also observed before [5]. For example, the mean of the reproduction of the pure rotational lines for HD<sup>16</sup>O is  $2.3 \times 10^{-5} \text{ cm}^{-1}$  if only the (0 0 0) transitions are employed in the MARVEL analysis. The agreement worsens to  $4.8 \times 10^{-5} \text{ cm}^{-1}$  if the full calibrated database is used to predict the pure rotational transitions.

Another source of inaccuracy when combining measured results from several sources is due to pressure shifts. This is especially of concern for this study as the pressure of HDO had to be increased several times in order to allow the detection of the HDO lines. Since there is no theory of pressure shifts of sufficient accuracy, no attempt was made in this work to correct the observed lines to, for example, zero pressure. This introduces a perhaps appreciable uncertainty for most of the measured lines. This is reflected in MARVEL uncertainties larger than otherwise expected for many rotational-vibrational levels.

Of all the transitions outside of the micro- and millimeterwave regions treated in this study, the most accurate one available is at 1480.094 038 033(67)  $\text{cm}^{-1}$  [76], observed for HD<sup>16</sup>O. Toth [72] has measured the same transition although with much less precision. Guelachvili [47] also measured this transition (as expected, his measured value only agrees with the other ones after the multiplicative recalibration of his results). As the present study of all the measured transitions of HD<sup>16</sup>O shows, the lower (0 0 0)  $5_{15}$  and the upper (0 1 0)  $5_{24}$  energy levels of this transition are involved in 461 and 76 transi-



tions, respectively. This means that MARVEL can fix the energies of these levels, and thus the transition, with considerable certainty. The MARVEL wave number and uncertainty of this transition is  $1\,480.094\,038\,033(7000)$   $\text{cm}^{-1}$ . Therefore, it is clear that MARVEL (a) did not change the value of the transition known accurately and precisely, and (b) increased the uncertainty of the measured line considerably. None of this is surprising given the large number of much less precisely determined transitions the two energy levels are involved in and the fact that MARVEL did not find it necessary to adjust the uncertainty of the original experimental measurement (which confirms its high precision).

Table 6: MARVEL vibrational band origins (VBO) for HD<sup>17</sup>O, with normal-mode ( $v_1v_2v_3$ ) assignments, MARVEL uncertainties, and the number of rotational levels (RL) the vibrational levels are holding within the present database.<sup>a</sup>

$v_1v_2v_3$	VBO/cm <sup>-1</sup>	Unc.	RL
0 0 0	0.000000 <sup>b</sup>	0	86
0 1 0	1399.674626	182	76

<sup>a</sup> See footnote a to Table 5.

<sup>b</sup> See footnote b to Table 5.

Table 7: MARVEL and RITZ vibrational band origins (VBO) for HD<sup>18</sup>O, with normal-mode ( $v_1v_2v_3$ ) assignments, uncertainties, and the number of rotational levels (RL) the vibrational levels are holding within the related databases.<sup>a</sup>

$v_1v_2v_3$	MARVEL			RITZ[114]	
	VBO/cm <sup>-1</sup>	Unc.	RL	VBO/cm <sup>-1</sup>	Unc.
0 0 0	0.000000 <sup>b</sup>	0	164	0.000000 <sup>b</sup>	0
0 1 0	1396.266365	176	156	1396.266540	299
1 0 0	2709.284316	178	125	2709.284528	311
0 2 0	2767.209640	220	85	2767.212120	696
0 0 1	3696.330134	345	149	3696.330305	696
1 1 0	4080.544716	385	99	4080.544949	984
0 3 0	4121.754251	349	77	4121.754430	696
0 1 1	5071.496323	352	114	5071.496500	696
0 4 0			7		
1 2 0			1		

2 0 0	5335.360667	485	76	5335.360899	984
1 0 1	6390.982474	189	132	6390.982322	334
0 2 1	6425.951205	175	109	6425.951439	342
2 1 0	6711.672686	173	80	6711.672989	342
0 0 2	7229.185072	227	134	7229.185309	389
0 5 0			12		
0 3 1			10		
1 1 1			11		
1 2 1			1		
3 0 0	7876.171252	399	46	7876.171489	676
0 2 2			23		
1 0 2	9930.723962	968	26	9930.724199	1634
0 1 2			79		
2 0 1			3		
0 0 3			61		
0 7 0			1		
3 1 0			5		
0 1 3			69		
4 0 0			4		
3 0 1			1		

<sup>a</sup> See footnote a to Table 5.

<sup>b</sup> See footnote b to Table 5.

Table 8: Comparison of HD<sup>16</sup>O, HD<sup>17</sup>O, and HD<sup>18</sup>O transition data found in HITRAN [141] and used in the present compilation.

	HD <sup>16</sup> O	HD <sup>17</sup> O	HD <sup>18</sup> O
Total number of transitions in present database	54 740	485	8 729
Unique transitions in present database	36 690	445	7 186
Number of validated transitions	53 291	478	8 634
Transitions in HITRAN database	13 238	175	1 611
Concordant transitions <sup>a</sup>	9 627	169	848
Transitions absent in present database	3 611	6	763
Transitions absent in HITRAN	27 063	276	6 338

<sup>a</sup> Transitions which are present both in HITRAN and in the present validated IUPAC database.

## 4. Discussion

### 4.1. A comparison with HITRAN

The relevant features of the HITRAN database [141, 142] are summarized in the original publications and in Part I; thus, they are not repeated here in detail, and just a few remarks are made. Table 8 compares the data used in this study with those in the HITRAN [141] database. There are several notable differences between HITRAN and the (present) IUPAC water-vapor data analysis. First, HITRAN contains a mixture of observed and calculated entries, whereas the lines used in the present analysis are based strictly on published experimental values. Even though the goal in HITRAN is to have a complete set of self-consistent and maintainable values, in practice this is

not possible at this time. Entries that have been determined under controlled laboratory conditions and are very accurate are frequently used in HITRAN. However, in order to have adequate sets of lines for important vibration-rotation bands, it is necessary to use theoretical extrapolations of the data (especially to the higher rotational levels), to have calculations of lines that are blended or obscured in laboratory observations, or to have calculations for quantities that theory actually yields better constrained values at present than experiment (some line-shape parameters, for example). Second, in the case of multiple sets of data for a quantity, HITRAN takes each datum from a single source. This consolidation of data has been made by judgment of the quality of the different data sets. With the superior structure of the IUPAC database, all high-quality data have been retained. In principle one can use a filtering code to produce a smaller list for a specific application. Thirdly, HITRAN is aimed at atmospheric temperature applications and, by application of a minimum transition intensity at 296 K requirement, rejects many hot transitions; our present analysis aims to capture all measured transitions. The increased number of transitions in the present compilation is partly a reflection of this.

The line positions of HD<sup>16</sup>O in the pure rotational band have remained unchanged in HITRAN since the 1986 edition [132] and originate from an older version of the JPL catalog [143]. It was generated based on the constants obtained in the simultaneous fit of submillimeter measurements of Messer *et al.* [49] and a set of experimentally determined ground state energy levels based on infrared measurements in the  $\nu_2$  band. All the lines from 500 to 7 514 cm<sup>-1</sup> originate from the SISAM database [144]. Most of the HD<sup>16</sup>O lines in SISAM are experimentally observed, while 1 525 were calculated based on the experimentally determined energy levels. There are no HD<sup>16</sup>O lines between 7 514 and 9 751 cm<sup>-1</sup>. From 9 751 to 10 782 cm<sup>-1</sup>, data from Tolchenov and Tennyson [107] were used, and starting from 11 750 cm<sup>-1</sup> lines were taken from Voronin *et al.* [100].

For HD<sup>18</sup>O, the current version of the JPL catalogue [145] was used in HITRAN for the pure rotational band and data from Toth's website [144] was used for all other lines from 500 to 3 825 cm<sup>-1</sup>. Note that line positions of 56 of these lines were determined from experimentally-derived energy levels rather than directly measured. The majority of 6 338 lines in the IUPAC dataset of experimental transitions that are absent in HITRAN (see Table 8) are lines above 3 825 cm<sup>-1</sup>.

For HD<sup>17</sup>O, the only lines in HITRAN are those from the SISAM database [144], while IUPAC data includes measurements in the pure rotational band [111, 110] and the  $\nu_2$  band from 93Tothb [61].

#### 4.2. *Completeness of data*

As shown in Part I, a comparison between the variational nuclear motion calculations and the results of the MARVEL analysis can be used to assess the overall completeness of the spectroscopic data for each water isotopologue and to indicate key regions for further experimental study.

Figures 4, 5, and 6 show comparisons of numbers of room temperature transitions for HD<sup>16</sup>O, HD<sup>17</sup>O, and HD<sup>18</sup>O, respectively. Each figure gives the measured transitions, the number of new transitions predicted from the MARVEL energy levels, the sum of these and the transitions predicted by the variational calculations. For simplicity the VTT [118] calculated transition intensities are assumed throughout.

Table 9: Proportion of the total integrated absorption<sup>a</sup> at 296 K, as estimated using the VTT [118] variational line list (“calc”), of the three isotopologues of HDO recovered experimentally (“expt”) and by including the extra transitions predicted by the MARVEL protocol (“MARVEL”). Note that different wavenumber maximum ( $\nu_{\max}$  in  $\text{cm}^{-1}$ ) and intensity cut-offs have been used for each isotopologue; these choices were dictated by the available experimental data.<sup>a</sup>

	$\nu_{\max}$	Cut-off	$A(\text{calc})$	$A(\text{expt})$	$\frac{A(\text{obs})}{A(\text{calc})}$	$A(\text{MARVEL})$	$\frac{A(\text{expt})+A(\text{MARVEL})}{A(\text{calc})}$
HD <sup>16</sup> O	25 000	10(−29)	5.95(−17)	5.13(−17)	86 %	8.66(−18)	100 %
HD <sup>17</sup> O	5 000	10(−28)	5.74(−17)	9.22(−18)	16 %	3.43(−17)	66 %
HD <sup>18</sup> O	20 000	10(−27)	5.79(−17)	3.50(−17)	60 %	1.60(−17)	88 %

<sup>a</sup> Absorption ( $A$ ) and intensities in  $\text{cm molecule}^{-1}$ ; powers of ten given in parenthesis.

One interesting measure of completeness is the proportion of the integrated absorption recovered for each isotopologue, a summary of which is given in Table 9. For HD<sup>16</sup>O, the measurements recover 86 % of the total absorbance at 296 K. The extra transitions involving energy levels predicted by MARVEL yield 14 % of the total absorbance meaning that the combined transitions lead to essentially total recovery of the absorbance. This result is significant because, given that the *ab initio* calculated transition intensities have been proven to be reliable [146, 147], it means very reliable estimates of missing absorption can be generated using very precise estimates of the frequencies. We note that at higher temperatures we expect the percentage of the total absorbance recovered experimentally to drop.

Unsurprisingly, the proportion of data observed for the minor isotopologues is less than for HD<sup>16</sup>O. Under these circumstances the use of MARVEL alone cannot provide all the missing absorption. However, improvement

in the simulation of the bands where only partial data is available is still significant.

## 5. Conclusions

Among many other applications of such data, a huge quantity of high-quality molecular data are needed for the understanding of spectroscopic measurements related to different stars and the atmospheres of planets and exoplanets. One must know precise positions and usually intensities and line shapes in order to be able to judge what is observed and how much of that species is present. Among the species for which spectroscopic data are needed for this or other important scientific or engineering applications, water is probably the single most important one. Thus, the study of the complete spectra of its isotopologues is of prime importance. At the same time, the high-resolution rovibrational spectra of the isotopologues of the water molecule form a fertile testing ground for different experimental and theoretical approaches yielding the required information.

While the ambitious aim of the IUPAC Task Group partially responsible for this work is to create a complete linelist for all isotopologues of water, a first step is to determine energy levels and line positions. This paper provides, for the first time, a dependable and carefully validated set of energy levels and transition wavenumbers, all with dependable and self-consistent uncertainties, for HD<sup>16</sup>O, HD<sup>17</sup>O, and HD<sup>18</sup>O. As proven in this work, the MARVEL approach [3, 4, 5], combined with results from variational nuclear motion computations, provides an ideal platform to achieve this goal.

Our analysis of the available HDO line frequency measurements pointed towards some issues with calibration of the published data. We were not always able to resolve these issues and it is our view that an effort to develop modern frequency standards, such as that of Guelachvili *et al* [130], is urgently needed.



The distributed information system W@DIS [148, 149], containing the transitions and the energy levels of this study, can be accessed via <http://wadis.saga.iao.ru>.

## Acknowledgements

We all thank the International Union of Pure and Applied Chemistry for funding under project 2004-035-1-100 (A database of water transitions from experiment and theory). In addition, this work has received partial support from the UK Engineering and Physical Science Research Council, the UK Natural Environment Research Council, the Royal Society, grant WWLC-008535-(reintegration) MCA FP6 EC, the Scientific Research Fund of Hungary (grant OTKA K77825), the European Union QUASAAR Marie Curie research training network, NATO, the Russian Foundation for Basic Research, the Belgian Federal Science Policy Office (contracts EV/35/3A, SD/AT/01A, PRODEX 1514901NLSFe(IC)), the Belgian National Fund for Scientific Research (FRFC contracts), the Communauté de Belgique (Action de Recherche Concertées), the NASA laboratory astrophysics program, NASA Earth Observing System (EOS), under grant NNX08AD92G, and the Programme National LEFE (CHAT) of CNRS (INSU). Part of the research described in this paper was performed at the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration

## Appendix A

The MARVEL protocol involves active maintenance (the A in MARVEL) and periodic update of the databases containing the measured transitions. Since the publication of Part I [1], new measurements [108, 150] came to our attention for  $\text{H}_2^{17}\text{O}$  and  $\text{H}_2^{18}\text{O}$ , the two water isotopologues treated there. Furthermore, during preparation of the Part II manuscript the MARVEL procedure, as noted above, has also been slightly adjusted in order to determine more reliable energy values and uncertainty estimates via a post-MARVEL validation scheme. Therefore, in this appendix we give improved VBOs for  $\text{H}_2^{17}\text{O}$  and  $\text{H}_2^{18}\text{O}$  and in the Supplementary Material we provide new or improved MARVEL energy levels for these two isotopologues, as well. These new data should replace the ones presented in Part I.

### A.1 $\text{H}_2^{17}\text{O}$

Two new sources have been added to the Part I database, 09LiNaKaCa [108] and 86GuRa [150]. Characteristics of these two sources are given in Table 10. Compared to Part I, the number of validated transitions increased from 8 614 to 9 028 and the number of energy levels increased from 2 687 to 2 723.

Application of the post-MARVEL validation scheme also meant that the MARVEL energy levels for this isotopologue changed as compared to Part I. The full set of VBOs and their improved uncertainties are given in Table 10. There are no new VBOs compared to Part I. The improved set of MARVEL rotational-vibrational energy levels are provided in the Supplementary Material.

Table 10: Data sources additional to those presented in Part I and their characteristics for H<sub>2</sub><sup>17</sup>O.<sup>a</sup>

Tag	Range (cm <sup>-1</sup> )	Trans. A/V	Physical conditions				Comments
			T(K)	p(hPa)	Abun.(%)	Rec.	
86GuRa [150]	100–1986	203/202					(10a)
09LiNaKaCa [108]	5907–6726	212/212	299	0.13-13	0.3	CRDS	(10b)

<sup>a</sup> See footnote a to Table 1.

(10a) 86GuRa MARVEL cuts the transition 83GuRa.5 at 207.73575 cm<sup>-1</sup>.

(10b) 09LiNaKaCa H<sub>2</sub><sup>17</sup>O transitions assigned in the CRDS spectrum of a highly <sup>18</sup>O enriched sample.

Table 11: MARVEL vibrational band origins (VBO), with polyad quantum number ( $P = 2v_1 + v_2 + 2v_3$ ), normal-mode ( $v_1v_2v_3$ ) and local-mode ( $(mn)^\pm v_2$ ) assignments, MARVEL uncertainties (Unc.), and the number of rotational levels (RL) the vibrational levels are holding within the present, improved database, for H<sub>2</sub><sup>17</sup>O.<sup>a</sup>

$P$	$v_1v_2v_3$	$(mn)^\pm v_2$	VBO/cm <sup>-1</sup>	Unc.	RL
0	0 0 0	(00) <sup>+</sup> 0	0.000000 <sup>b</sup>	0	192
1	0 1 0	(00) <sup>+</sup> 1	1591.325719	47	153
2	0 2 0	(00) <sup>+</sup> 2	3144.980520	46	78
2	1 0 0	(10) <sup>+</sup> 0	3653.142275	28	108
2	0 0 1	(10) <sup>-</sup> 0	3748.318105	31	141
3	1 1 0	(10) <sup>+</sup> 1	5227.705615	473	69
3	0 1 1	(10) <sup>-</sup> 1	5320.251393	527	157
4	1 2 0	(10) <sup>+</sup> 2	6764.725615	473	73
4	0 2 1	(10) <sup>-</sup> 2	6857.272485	47	94
4	2 0 0	(20) <sup>+</sup> 0	7193.246415	47	81
4	1 0 1	(20) <sup>-</sup> 0	7238.714012	51	110
4	0 0 2	(11) <sup>+</sup> 0	7431.076115	473	27
5	1 1 1	(20) <sup>-</sup> 1	8792.544010	360	108
5	0 1 2	(11) <sup>+</sup> 1	8982.869215	473	55
6	1 2 1	(20) <sup>-</sup> 2	10311.201787	440	75
6	2 0 1	(30) <sup>-</sup> 0	10598.475620	492	103
6	1 0 2	(21) <sup>+</sup> 0	10853.505315	473	69
6	0 0 3	(21) <sup>-</sup> 0	11011.882910	2465	84

7	1 3 1	(20) <sup>-3</sup>	11792.822200	490	40
7	2 1 1	(30) <sup>-1</sup>	12132.992610	493	94
7	0 1 3	(21) <sup>-1</sup>	12541.227060	517	51
8	2 2 1	(30) <sup>-2</sup>	13631.499810	542	52
9	0 7 1	(10) <sup>-7</sup>	13808.273310	493	2
8	3 0 1	(40) <sup>-0</sup>	13812.158110	493	74
8	1 0 3	(31) <sup>-0</sup>	14296.279510	197	36
10	3 2 1	(40) <sup>-2</sup>	16797.167510	3944	70
10	4 0 1	(50) <sup>-0</sup>	16875.620510	3944	55

<sup>a</sup> The uncertainties (Unc.) are given in units of  $10^{-6}$   $\text{cm}^{-1}$ . The VBOs are ordered by energy.

<sup>b</sup> The value of the vibrational ground state was fixed to zero with zero uncertainty.

## A.2 $\text{H}_2^{18}\text{O}$

Three new sources have been added to the Part I database, 09LiNaKaCa [108], 09PuCaHaVa [151], and 86GuRa [150]. Characteristics of these sources are given in Table 12. There are 6 orphans in the database and an FSN containing 22 transitions. Compared to Part I, the number of validated transitions increased from 29 364 to 31 705 and the number of energy levels increased from 4 849 to 5 131.

Application of the post-MARVEL validation scheme also meant that the MARVEL energy levels for  $\text{H}_2^{18}\text{O}$  isotopologue changed as compared to Part I. The full set of VBOs and their improved uncertainties are given in Table 12. The improved set of MARVEL energy levels are provided in the Supplementary Material.

Table 12: Data sources additional to those presented in Part I and their characteristics for  $\text{H}_2^{18}\text{O}$ .<sup>a</sup>

Tag	Range ( $\text{cm}^{-1}$ )	Trans. A/V	Physical conditions				Comments
			$T(\text{K})$	$p(\text{hPa})$	Abun.(%)	Rec. $L(\text{m})$	
09PuCaHaVa [151]	6.4–18.4	5/5	RT				
86GuRa [150]	55 – 2 053	351/350					(12a)
09LiNaKaCa [108]	5 907–6 726	2015/2006	299	0.13-13	84-92	CRDS	

<sup>a</sup> See footnote a to Table 2.

(12a) 86GuRa MARVEL cuts the transition 83GuRa.21 at  $147.193\ 80\ \text{cm}^{-1}$ .

Table 13: MARVEL vibrational band origins (VBO), with polyad quantum number ( $P = 2v_1 + v_2 + 2v_3$ ), normal-mode ( $v_1v_2v_3$ ) and local-mode ( $(mn)^\pm v_2$ ) assignments, MARVEL uncertainties, and the number of rotational levels (RL) the vibrational levels are holding within the present revised database, for  $\text{H}_2^{18}\text{O}$ .<sup>a</sup>

$P$	$v_1v_2v_3$	$(mn)^\pm v_2$	VBO/ $\text{cm}^{-1}$	Unc.	RL
0	0 0 0	(00)+0	0.000000	0	298
1	0 1 0	(00)+1	1588.276029	22	227
2	0 2 0	(00)+2	3139.050028	36	146
2	1 0 0	(10)+0	3649.685417	50	192
2	0 0 1	(10)-0	3741.566775	45	216
3	0 3 0	(00)+3	4648.477809	341	130
3	1 1 0	(10)+1	5221.243308	377	170
3	0 1 1	(10)-1	5310.461383	281	246
4	1 2 0	(10)+2	6755.510366	283	167
4	0 2 1	(10)-2	6844.598785	194	169
4	2 0 0	(20)+0	7185.878096	506	176
4	1 0 1	(20)-0	7228.883501	493	181
4	0 0 2	(11)+0	7418.724061	294	147
5	1 3 0	(10)+3	8249.037152	459	79
5	0 3 1	(10)-3	8341.107933	417	120
5	2 1 0	(20)+1	8739.525858	503	84
5	1 1 1	(20)-1	8779.719569	349	143
5	0 1 2	(11)+1	8967.565058	503	87
6	0 4 1	(10)-4	9795.331502	473	72
6	2 2 0	(20)+2	10256.584858	503	61

6	1 2 1	(20) <sup>-2</sup>	10295.634509	450	108
6	0 2 2	(11) <sup>+2</sup>	10483.221458	503	53
6	3 0 0	(30) <sup>+0</sup>	10573.916858	503	97
6	2 0 1	(30) <sup>-0</sup>	10585.285174	460	133
6	1 0 2	(21) <sup>+0</sup>	10839.955723	499	90
6	0 0 3	(21) <sup>-0</sup>	10993.681009	450	103
7	2 3 0	(20) <sup>+3</sup>	11734.525057	503	68
7	1 3 1	(20) <sup>-3</sup>	11774.707602	373	103
7	0 3 2	(11) <sup>+3</sup>	11963.537158	503	73
7	3 1 0	(30) <sup>+1</sup>	12106.977658	503	98
7	2 1 1	(30) <sup>-1</sup>	12116.797008	466	133
7	1 1 2	(21) <sup>+1</sup>	12372.705608	410	103
7	0 1 3	(21) <sup>-1</sup>	12520.122915	47	110
8	2 2 1	(30) <sup>-2</sup>	13612.710202	47	74
8	3 0 1	(40) <sup>-0</sup>	13795.398204	47	94
8	2 0 2	(31) <sup>+0</sup>	14187.982458	1358	49
8	1 0 3	(31) <sup>-0</sup>	14276.336099	464	68
10	3 2 1	(40) <sup>-2</sup>	16775.380902	1418	28
10	4 0 1	(50) <sup>-0</sup>	16854.990902	1418	44

<sup>a</sup> See footnotes a and b to Table 11.

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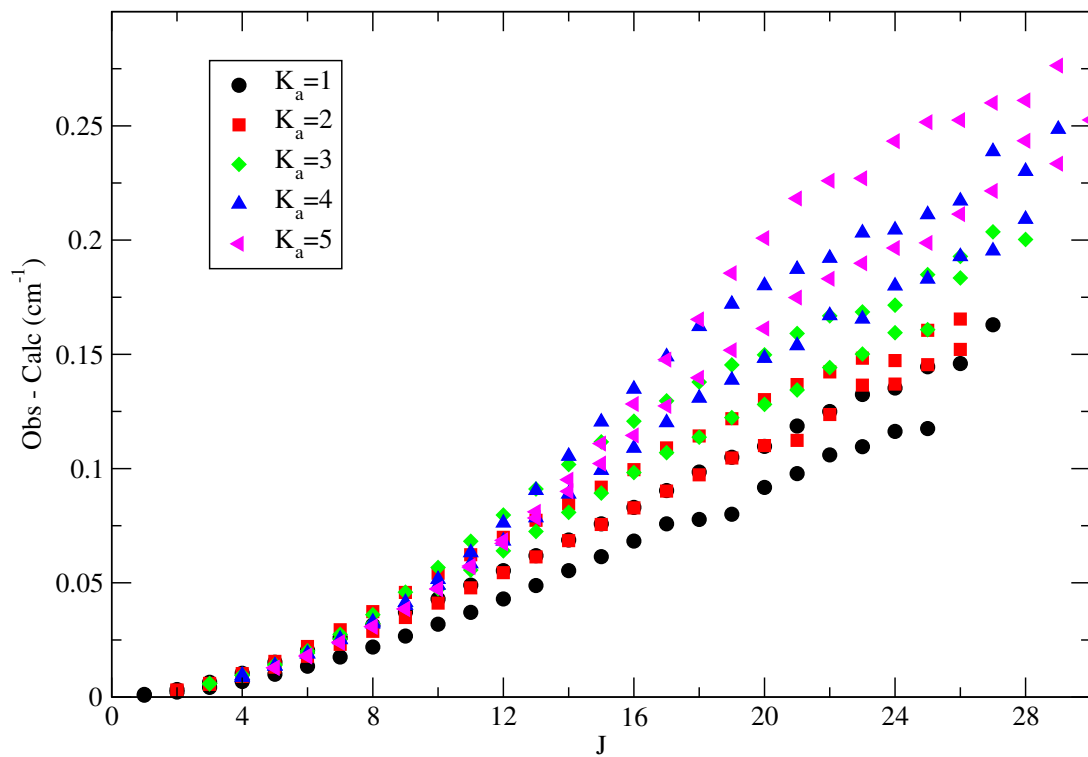


Figure 1: Residuals of “Observed” (MARVEL) minus calculated (VTT [118]) energy levels of HD<sup>16</sup>O for the ground vibrational state and  $K_a$  between 1 and 5.

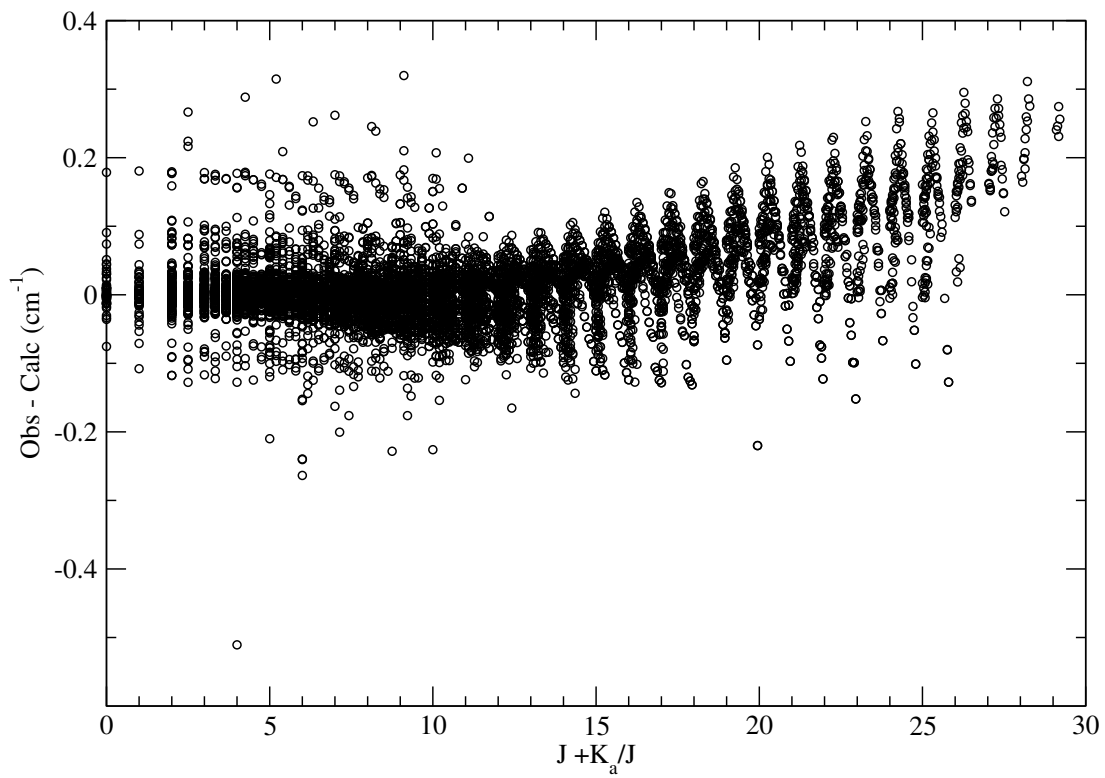


Figure 2: Residuals of “Observed” (MARVEL) minus calculated (VTT [118]) for all energy levels of HD<sup>16</sup>O derived in this work.

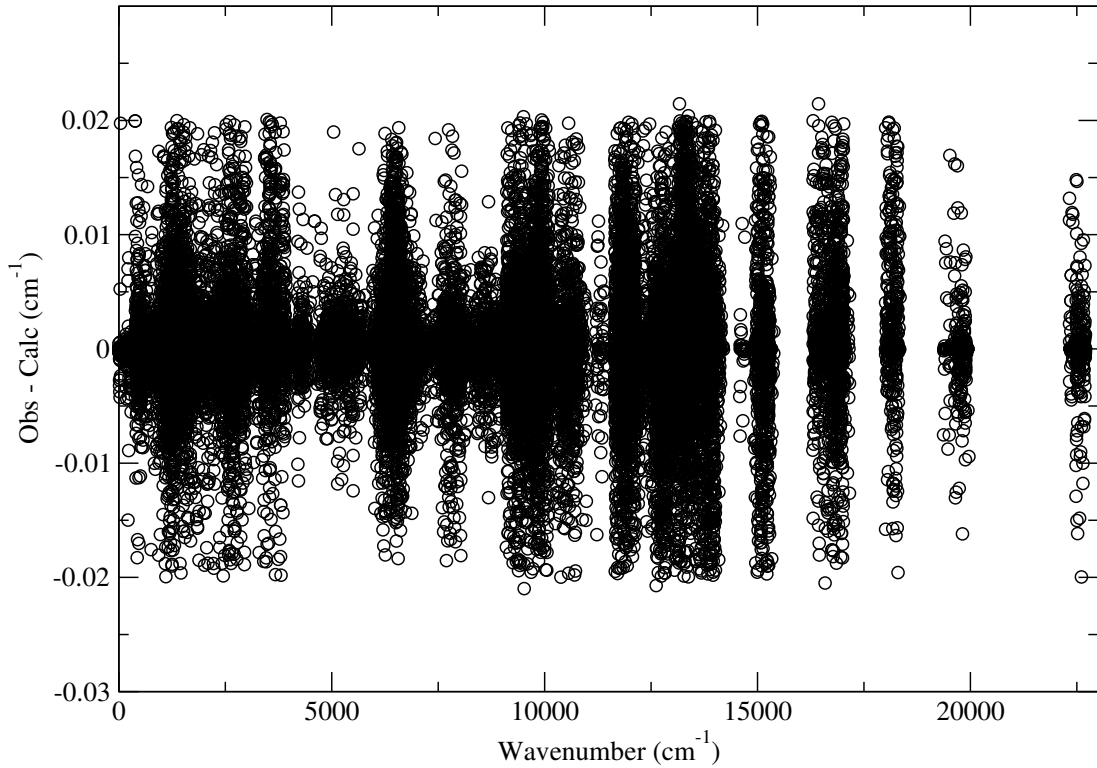


Figure 3: Differences between the measured transition wavenumbers and those calculated using the MARVEL energy levels derived in this work for HD<sup>16</sup>O.

Figure 4: Observed (top panel), MARVEL predicted (second panel), their sum (third panel) transitions based on predictions from the VTT variational calculations [118] (bottom panel) for HD<sup>16</sup>O. Calculated line intensities are used in each case.

Figure 5: Observed (top panel), MARVEL predicted (second panel), their sum (third panel) transitions based on predictions from variational calculations [116, 140] (bottom panel) for HD<sup>17</sup>O. Calculated line intensities are used in each case.

Figure 6: Observed (top panel), MARVEL predicted (second panel), their sum (third panel) transitions based on predictions from variational calculations [116, 140] (bottom panel) for HD<sup>18</sup>O. Calculated line intensities are used in each case.