J-PAS: Forecasts for dark matter - dark energy elastic couplings

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Abstract. We consider a cosmological model where dark matter and dark energy feature a coupling that only affects their momentum transfer in the corresponding Euler equations. We perform a fit to cosmological observables and confirm previous findings within these scenarios that favour the presence of a coupling at more than 3σ. This improvement is mainly driven by cluster counts from Planck Sunyaev-Zeldovich data that we include as a certain prior. We subsequently perform a forecast for future J-PAS data and find that clustering measurements will permit to clearly discern the presence of an interaction within a few percent level with the uncoupled case at more than 10σ when the complete survey, covering 8500 sq. deg., is considered. We found that the inclusion of weak lensing measurements will not help to further constrain the coupling parameter. For completeness, we compare to forecasts for DESI and Euclid, which provide similar discriminating power.

Contents

1 Introduction

The upcoming generation of cosmological observational campaigns will crucially impact our understanding of the universe and, in particular, the physics of its dark sector. Among the available data in the near future, the on-going and planned galaxy surveys will contribute a very rich source of information able to put stringent constraints on the statistical properties of the large scale structures in the universe. Three different kinds of galaxy surveys are currently operating according to the method used for the determination of redshifts. On one hand, spectroscopic surveys that can perform high-precision redshifts measurements from high-quality spectra of a pre-selected sample of galaxies. Examples include the current BOSS [1] and the future spectroscopic Euclid [2] and DESI [3]. Second, photometric surveys that obtain photo-spectra using photometry with a reduced number of filters. These surveys can build larger catalogues of objects, but with poorer redshift accuracies than the spectroscopic

ones. The best example of a current photometric survey is DES [4] and among the future ones we have the photometric Euclid survey [2] and LSST [5]. Finally, there is a third class of surveys, the so-called spectro-photometric surveys which produce high-quality pseudospectra by combining photometry in many different frequencies obtained through the combination of broad, medium and narrow band filters. From those pseudospectra, precise redshifts can be obtained for a high number of sources. Among the current and future spectro-photometric surveys we have PAU [6], J-PLUS [7] and the future J-PAS [8].

A very relevant question that galaxy surveys will be capable of unveiling is the existence and nature of a coupling among the dark components. In this respect, many possibilities have been explored in the literature [9–19], most of which modify the background evolution of the universe and, consequently, they are already constrained by geometrical tests. In this work, we are interested in another class of couplings with the remarkable property of leaving the background evolution unmodified and only affecting the perturbations sector. In particular, they impact the clustering of dark matter so that the aforementioned galaxy surveys will provide a crucial tool to test these couplings by means of confronting the inferred statistical properties of the dark matter distribution with the collected data. These couplings have been considered in the literature arising from different scenarios, e.g., an elastic Thomson-like scattering between dark matter and dark energy [20–23], from a fluid description of dark matter interacting with dark energy (dubbed pure momentum exchange) [24–30] or, from a more phenomenological perspective, as a coupling proportional to the relative velocities of the dark components [31]. In all the cases, the background and the continuity equations of the perturbations are not modified and only the Euler equations of dark matter and dark energy exhibit new terms driven by the coupling. It is remarkable that these couplings provide promising prospects to alleviate the existing tension in the amplitude of the dark matter clustering as measured from low redshift data [32–35] and the one inferred from CMB observations [36].

In this work, we will focus on the class of models where dark matter and dark energy feature a coupling proportional to their relative velocities that will also serve as proxy for more general scenarios where similar modified equations arise. By the very construction of the coupling and by virtue of the Cosmological Principle, the background evolution is unaffected and only the Euler equations of the perturbations are modified at first order in cosmological perturbations. For this model we will perform a fit to current cosmological observables to confirm previous findings with updated data, namely that a non-vanishing coupling parameter is favoured by data, with the uncoupled case at more than 3σ. This result is only obtained when including a prior based on cluster counts from Sunyaev-Zeldovich data. We will discuss the consistency of including this data in the form of a prior. In any case, this intriguing result motivates to carry out a forecast with future galaxy surveys to discern to which level of accuracy they will help to discriminate the existence of this coupling in the dark sector. As we will show, they will indeed play a crucial role to probe the existence of a coupling in the dark sector.

2   Elastic couplings in the dark sector

In this section we will introduce the relevant equations that govern the dynamics of the model under consideration and define its relevant parameters. We will not enter into the details of the model and we refer to [31] for a more comprehensive discussion.

We will assume that the content of the universe can be described by a stress energy tensor of the perfect fluid form

$$T^{\mu\nu} = \sum_i \left[(\rho_i + p_i) u_i^\mu u_i^\nu + p_i g^{\mu\nu} \right] , \tag{2.1}$$

where the sum runs over all the components in the universe, namely baryons, photons and neutrinos together with a dark sector that we assume conformed by cold dark matter and dark energy with a constant equation of state parameter w. While the ordinary matter sector satisfies the usual conservation laws $\nabla_\mu T^{\mu\nu} = 0$, the dark sector will be assumed to satisfy the non-conservation equations

$$\nabla_\mu T_{\text{dm}}^{\mu\nu} = Q^\nu , \tag{2.2}$$
$$\nabla_\mu T_{\text{de}}^{\mu\nu} = -Q^\nu , \tag{2.3}$$

where Q^ν describes the interaction that is proportional to the relative velocity as follows

$$Q^\nu = \bar{\alpha} \left(u_{\text{de}}^\nu - u_{\text{dm}}^\nu \right) . \tag{2.4}$$

The parameter $\bar{\alpha}$ measures the strength of the interaction and can depend both on time and space, although, for the sake of simplicity, throughout this work we will mostly assume it constant. From the form of the coupling, it is clear that the background equations are not modified because both dark components have the same background 4-velocities. The perturbed metric in the Newtonian gauge reads

$$\mathrm{d}s^2 = a^2(\tau) \left[-(1 + 2\Psi)\mathrm{d}\tau^2 + (1 - 2\Phi)\mathrm{d}\vec{x}^{\,2} \right] , \tag{2.5}$$

and the equations governing the dark sector perturbations, in the absence of anisotropic stress so that $\Phi = \Psi$, are the following:

$$\delta_{\text{dm}}' = -\theta_{\text{dm}} + 3\Phi' , \tag{2.6}$$
$$\theta_{\text{dm}}' = -\mathcal{H}\theta_{\text{dm}} + k^2\Phi + \Gamma(\theta_{\text{de}} - \theta_{\text{dm}}) , \tag{2.7}$$
$$\delta_{\text{de}}' = -3\mathcal{H} \left(c_{\text{s}}^2 - w \right) \delta_{\text{de}} + 3(1 + w)\Phi' - \theta_{\text{de}}(1 + w) \left(1 + 9\mathcal{H}^2 \frac{c_{\text{s}}^2 - w}{k^2} \right) , \tag{2.8}$$
$$\theta_{\text{de}}' = \left(-1 + 3c_{\text{s}}^2 \right) \mathcal{H}\theta_{\text{de}} + k^2\Phi + \frac{k^2 c_{\text{s}}^2}{1 + w}\delta_{\text{de}} - \Gamma R(\theta_{\text{de}} - \theta_{\text{dm}}) , \tag{2.9}$$

where we have defined the usual density contrast $\delta \equiv \delta\rho/\rho$ and Fourier velocity perturbation $\theta \equiv i\vec{k} \cdot \vec{v}$. Throughout this work, the dark energy sound speed will be assumed to be $c_{\text{s}} = 1$. Moreover, we have introduced the quantities

$$\Gamma \equiv \bar{\alpha}\frac{a}{\rho_{\text{dm}}} , \tag{2.10}$$
$$R \equiv \frac{\rho_{\text{dm}}}{(1 + w)\rho_{\text{de}}} . \tag{2.11}$$

The quantity Γ can be interpreted as the interaction rate between dark matter and dark energy. Thus, the relative hierarchy between Γ and the Hubble expansion rate $\mathcal{H}$ will determine the relevance of the coupling. For a constant $\bar{\alpha}$, we can see that $\Gamma \propto a^4$ so that it grows as the universe expands. Hence, the coupling is irrelevant at sufficiently high redshift

and it only affects the evolution of the perturbations at low redshift when $\Gamma \gtrsim \mathcal{H}$. On the other hand, the dark matter-to-dark energy ratio R measures the relative importance of the coupling for dark matter and dark energy according to their relative abundances. Since this quantity can be written as $R = \frac{\Omega_{\rm DM}}{(1+w)\Omega_{\rm DE}} a^{3w}$, we can see that it is large throughout the universe evolution so that the coupling is larger in the dark energy equations. However, it is the effect on the dark matter perturbations that will impact their evolution (see [31]). Finally, since the coupling parameter $\bar{\alpha}$ has mass dimension 5, it is convenient to normalise it appropriately and work with the dimensionless parameter α defined as

$$\alpha \equiv \bar{\alpha} \frac{8\pi G}{3H_0^3} \ . \tag{2.12}$$

As we will see below, the normalisation has been chosen so that its natural value is $\mathcal{O}(1)$.

After this brief review of the theoretical model, we proceed to numerically analyse the evolution of the perturbations in the next Section.

3   Numerical analysis

In order to gain some intuition on the observable effects of the coupling, we have solved the system of equations (2.6)-(2.9) numerically. To do so we have modified the publicly available code CLASS [37] for the computation of the linear perturbations, using the Newtonian gauge. As explained in Section 2, the background cosmology remains unaffected and, therefore, we only need to modify the code by adding the new interacting term in the Euler equations, for both DE and DM. As a consequence, any deviation from the standard model is only due to the modification in the perturbation sector.

The model has also been explored modifying the publicly available code CAMB [38], in which the evolution of the cosmological perturbations is computed in the synchronous gauge. The comparison of gauge-invariant quantities then ensures the consistency of both numerical implementations. It is worth mentioning, however, that in the case of CAMB some extra considerations must be taken into account to correctly include the new interaction. Although only modifications in the velocity equations of the dark sector are needed also in the synchronous gauge (see Appendix A), using the residual gauge freedom of this gauge to set $\theta_{\rm dm} = 0$ at all times is no longer possible, precisely because the interaction sources the evolution equation for $\theta_{\rm dm}$. Therefore, a new evolution variable must be introduced to deal with the DM velocity evolution. A potential problem with this approach is that the gauge remains partially unfixed so we could obtain unphysical gauge solutions. This could be avoided by enlarging the content of the universe with a marginal dust component that can be used to completely fix the gauge (this approach has been used in e.g. [39]). We have corroborated however that the obtained numerical solutions correspond to physical modes and any potential gauge mode remains subdominant.

We will fix a fiducial model for our analysis in this Sec. with the cosmological parameters to $H_0 = 67.4\,\mathrm{km/s/Mpc}$, $\Omega_{\rm b} h^2 = 0.0224$, $\Omega_{\rm dm} h^2 = 0.120$, $w = -0.98$ and $c_{\rm s}^2 = 1$. We choose as reference model a wCDM rather than ΛCDM to clearly isolate the effects of the interaction and to avoid any degeneracies with the DE equation of state resulting in a cosmological constant or not, as the interaction requires $w \neq -1$ and the cosmological constant does not have perturbations. Any other parameter is set to the default value of each code.

Before proceeding to the discussion of the results obtained from the modified numerical codes, it is convenient to warn that we are only considering the linear evolution and our

statements in the following refer to that regime. The gravitational collapse induces non-negligible non-linear effects on small scales and the scale at which non-linearities become important shifts towards larger scales throughout the universe evolution so the linear regime becomes valid up to smaller values of k at lower redshift. This non-linear evolution will modify our findings so our statements only hold up to the scale of non-linearity. Beyond that scale, we should resort to proper N-body simulations to draw reliable conclusions. Simulations for a similar model were performed in e.g. [21, 22], although we find that the different time-evolution of the coupling constant casts doubts on the possibility of making a direct translation of their results to our case.

3.1 Constant coupling

Let us start by revisiting the scenario studied in Ref. [31], where the coupling parameter has neither time nor scale dependence: $\alpha(\tau, k) = \alpha$. In the following, we will perform a somehow more extensive analysis of the perturbations evolution from the numerical solutions obtained by using the aforementioned cosmological parameters and study the dependence on the coupling α.

3.1.1 Matter and CMB power spectra

The elastic coupling is driven by the relative velocity between DM and DE (eq. (2.7) and eq. (2.9), respectively), therefore, distinctive behaviours with respect to a non interacting model are expected to appear on small scales ($k \gg \mathcal{H}$), since for large scales both DM and DE share the same rest frame. Also, it is important to notice that the background does not differ from the uncoupled wCDM model, hence, all the effects are genuinely due to the modified perturbation sector.

In Figure 1, we show the matter power spectrum for different values of the coupling parameter α as well as its relative ratio with respect to the wCDM model for the same cosmological parameters. As commented above, the interaction leaves no imprint on large scales and only the small scales are affected where we see a suppression of power. This suppression exhibits two different regimes, namely: while for intermediate scales it shows a k-dependence, when we go to smaller scales it saturates becoming k-independent. The suppression in the matter power spectrum is larger as the value of the coupling parameter α is increased. A remarkable property of the coupling is its ability to shift the turnover of the matter power spectrum without modifying the matter-radiation equality time, i.e., for fixed matter and radiation density parameters. Notice the resemblance of the suppression to the case of massive neutrinos, which however does not modify the turnover for fixed radiation abundance. In the next Section 3.1.2, we will understand these features by looking at the evolution of the perturbations.

The effects of the elastic coupling on the CMB power spectra for temperature, polarisation and cross-correlations are shown in Figure 2 for different values of the parameter α, together with the current available CMB temperature and polarisation data from the Planck Collaboration [40], the Atacama Cosmology Telescope [41] and the South Pole Telescope [42]. The power spectrum for temperature is most significantly modified at large scales through the Integrated Sachs–Wolfe (ISW) effect, as expected because the interaction is relevant at very late times. The polarisation power spectra are however mainly modified at small scales increasing the amplitude of the high-ℓ oscillations due to lensing effects (these oscillations in the high-ℓ sector are also visible in the temperature power spectrum) and a non-oscillating
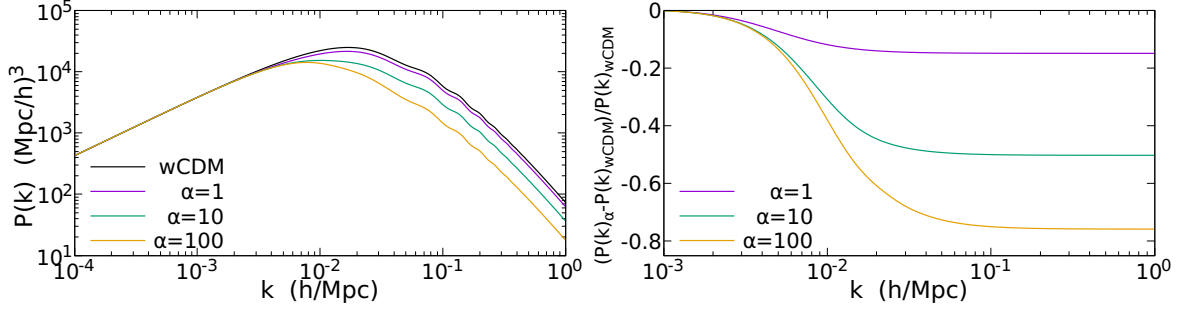
Figure 1: In the left plot, we show the matter power spectrum for several values of the coupling parameter α and for a wCDM model with the same cosmological parameters. In the right plot, we show the relative ratio of the matter power spectrum using different values of the coupling parameter α.
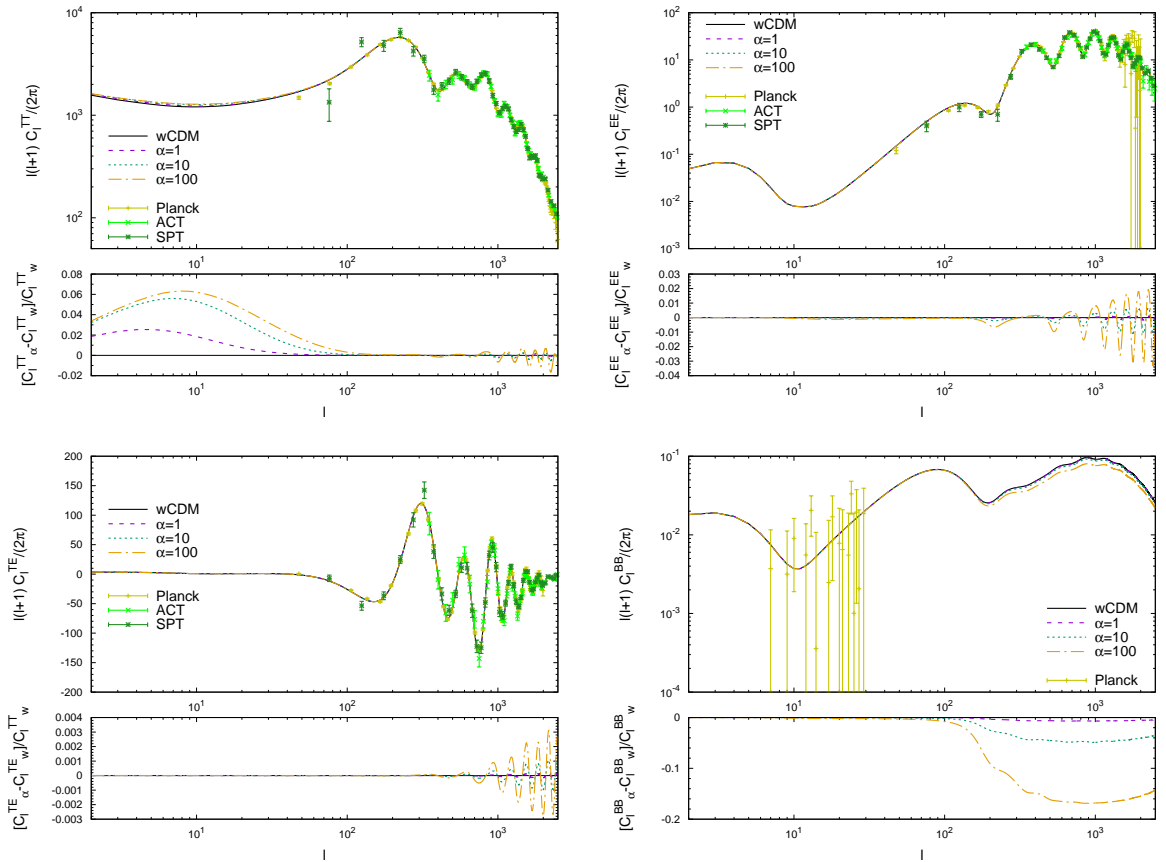
Figure 2: The CMB angular power spectra are presented for different values of the parameter α. The normalisation of relative errors is presented in all cases with respect to a wCDM model ($\alpha = 0$), and temperature and polarisation data from Planck [40], ACT [41] and SPT [42] are also included.

correction in the BB power spectrum. Since the coupling reduces the clustering of matter, the lensing potential decreases as α increases, as shown in Figure 3.
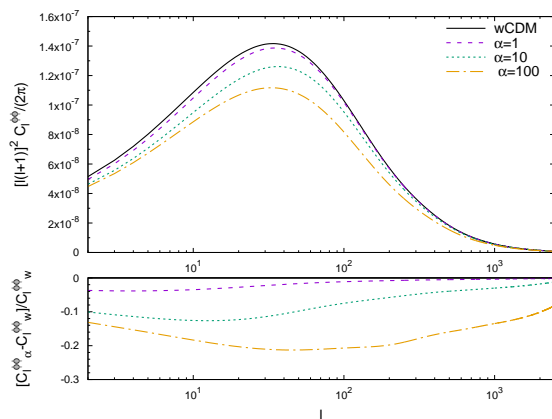
Figure 3: The amplitude of the lensing potential is sensitive to the parameter α, leading to modifications on the high-ℓ oscillations of both the temperature and polarisation power spectra.

### 3.1.2 Growth of structures: lowering the value of σ_8

When the interaction is efficient, DM and DE become tightly coupled so they tend to form a locked system, thus preventing the clustering of DM due to the DE pressure that causes the suppression in the matter power spectrum explained above. In order to gain a more thorough understanding of the underlying mechanism, we will analyse how the growth of structures is modified due to the coupling. With this in mind, this section is devoted to the study of the DM density contrast evolution as well as the Newtonian potential Φ. We will then explore the impact on the parameter σ_8, which will be relevant for the fit to data carried out in Section 4.

In Figure 4, we display the evolution of $\delta_{\rm dm}$ and Φ when the elastic coupling is efficient, that is, when $\Gamma \gg \mathcal{H}$. When the interaction dominates, at low redshift and small scales, the DM density contrast ceases growing and it saturates to a constant value. Concerning the Newtonian potential, the coupling induces an additional decrease, which will affect the late ISW effect. The freezing of the DM density contrast explains the suppression on the matter power spectrum shown in Figure 1 and it also explains why the gravitational potential exhibits a more pronounced decrease at late times.

A direct consequence of the explained features is that the amplitude of matter fluctuations is reduced. This can be clearly seen in Figure 5, where we plot the evolution of the ratio between the values of σ_8 for the coupled and uncoupled cases, keeping fixed the remaining cosmological parameters. Since the available amount of DM is limited, the suppression in σ_8 saturates to $\sim 30\%$ of the non interacting value as we consider very large values for α. The very same reason explains why, in the analogous case of interactions with baryons, this saturation is smaller as shown in [43].

3.1.3 Velocities evolution: new peculiar ones and Dark Acoustic Oscillations

The modified Euler equations of DM and DE induce an additional dragging on the DM component sourced by the DE pressure (let's recall that we are assuming $c_{\rm s}^2 = 1$) that modifies the evolution of the peculiar velocities. Furthermore, the form of the coupling, similar to the Thomson scattering of baryons and photons before recombination, leads to analogous
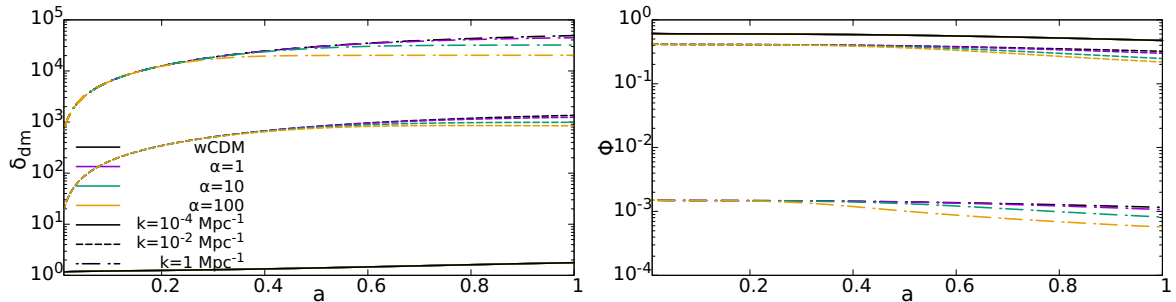
Figure 4: Evolution with the scale factor of the DM density contrast $\delta_{\rm dm}$ (left) and the Newtonian potential Φ (right) for different values of the coupling constant α and different modes. For these plots, we have used the Newtonian gauge. The black line represents the reference wCDM model while the purple, green and yellow lines are the elastic coupling model with $\alpha = 1, 10$ and 100 respectively. Solid lines represents the mode $k = 10^{-4}\,\mathrm{Mpc}^{-1}$, dashed lines the mode $k = 10^{-2}\,\mathrm{Mpc}^{-1}$ and dash-dotted lines the mode $k = 1\,\mathrm{Mpc}^{-1}$.
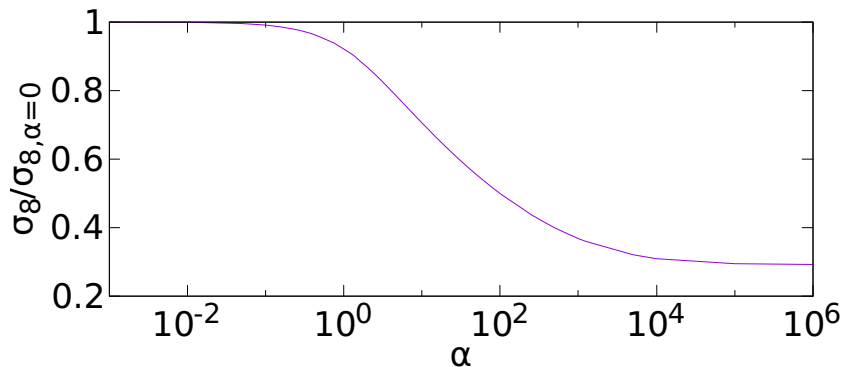
Figure 5: Evolution of the ratio of the value of σ_8 compared to a non interacting model value $\sigma_{8,\alpha=0}$ using the same cosmological parameters, for different values of the coupling parameter α.

oscillations in the dark sector that we refer to as Dark Acoustic Oscillations (DAO). These oscillations are more clearly visible in the evolution of the velocities, as shown in Figures 7 and 8.

In the left plot of Figure 7 we can see how the relative velocity of DM and DE decreases as the coupling α becomes larger, indicating that the larger α the more efficient is the tight coupling of these dark components. We also see how the oscillations are more noticeable for stronger couplings. These oscillations are also manifest in the spectra of the individual velocities shown in Figure 8. We can identify the three relevant regimes of the considered scenario. At very large scales, there are no deviations with respect to the wCDM model. As we consider smaller scales, the dragging that DE exerts on DM gives rise to the tight coupling that forces both dark components to move coherently and produces the aforementioned oscillations. At even smaller scales, DM tends to fall into the deeper gravitational wells and the DE dragging is not sufficient to prevent it, so both components no longer move coherently. However, the DE pressure still has an effect and delays the falling of DM into the wells, thus causing the appearance of additional peculiar velocities between DM and
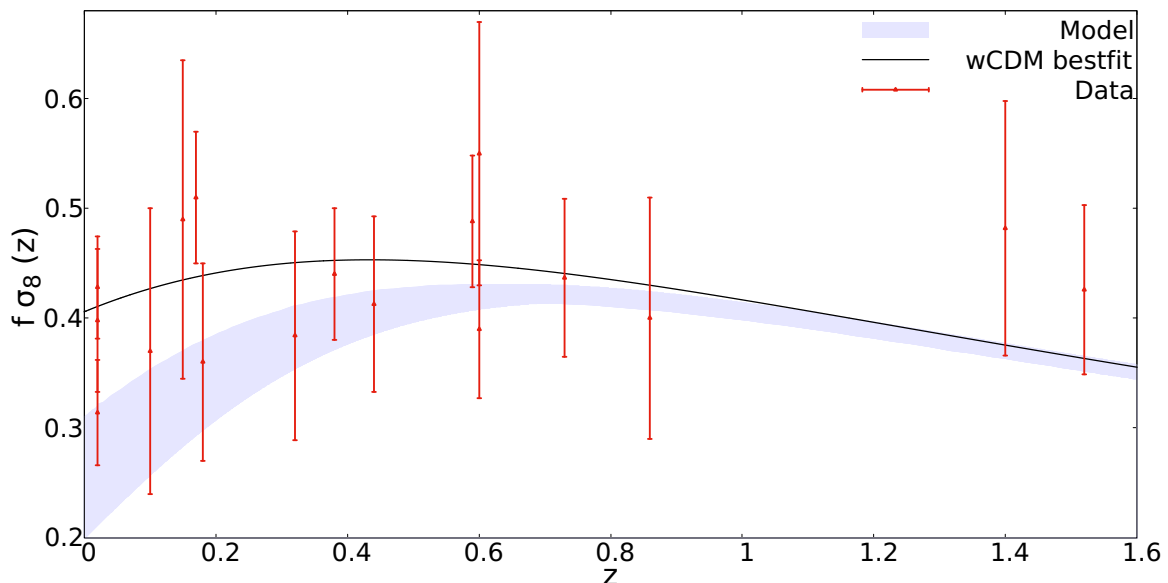
Figure 6: Evolution of $f\sigma_8(z)$ for the reference model wCDM (black line), for the interacting model (blue area) and data in Ref. [44] (red). The blue area encloses $f\sigma_8(z)$ evolution for values of the coupling parameter α inside the 1σ region of the fit of Table 2 $\alpha \in [0.675, 1.265]$, and for the different modes $k \in [0.007, 1]$ $h/$Mpc in order to take into account the scale dependence of $f\sigma_8(z)$ for this model. The other cosmological parameters are set to the mean value of the fit of Table 2, while the derived parameters are calculated by the code in each case. The data points are only for an illustrative goal, as they are obtained for several scales but assuming a k-independent model, while in our case $f\sigma_8(z)$ is k-dependent. Thus, before drawing any conclusion from the fit to these data, a more appropriate re-compilation of the points by including the scale-dependence of the model would be necessary.

baryons. As we can see, this regime occurs at smaller scales as we increase the value of α.[1] In this respect, we could think that these velocities would induce an additional velocity bias between the DM and baryons. However, we need to keep in mind that when we use galaxies to trace the underlying DM density field, these correspond to virialised objects within the associated DM haloes. For sufficiently small values of the interaction parameter α, we expect the gravitational bound of galaxies to the host DM halo to be sufficiently efficient as to make the galaxies follow the motion of the halo. In other words, galaxies are still good tracers of the DM distributions without additional velocity bias. Let us emphasise however once again that a full account of these effects would require resorting to a proper N-body simulation.

3.2 Screening on small scales

The elastic coupling model naturally prevents effects on very large scales as the relative velocities between the universe components are negligible. In this section, we want to explore the possibility of having a screening mechanism suppressing the interaction at some small k_s scale below the horizon. We will model this behaviour in a phenomenological manner by

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[1]It might be convenient to stress again that our analysis is limited to the linear regime. The non-linear evolution of the dark matter clustering in the presence of the friction with dark energy would favour the dissipation of angular momentum that could in turn speed up the clustering.
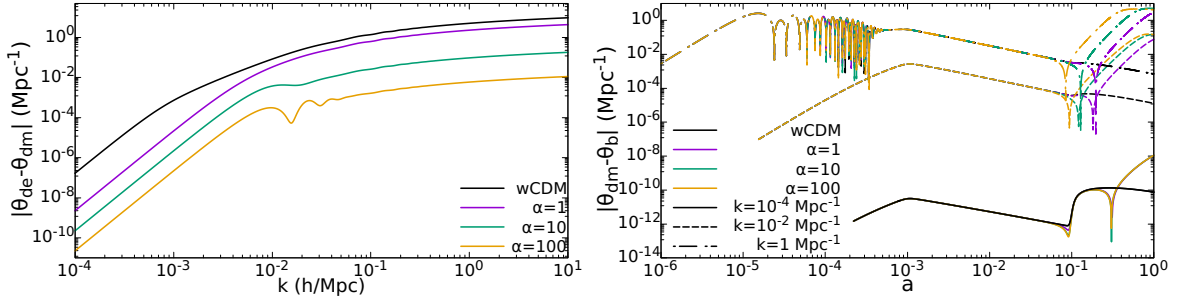
Figure 7: Relative velocity today between DE and DM today for different scales with and without interaction (left) and evolution of the relative velocity between matter components for different values of the coupling constant α and different modes (right). The black line represents the reference wCDM model while the purple, green and yellow lines are the elastic coupling model with $\alpha = 1, 10$ and 100 respectively. Solid lines represents the mode $k = 10^{-4}\,\mathrm{Mpc}^{-1}$, dashed lines the mode $k = 10^{-2}\,\mathrm{Mpc}^{-1}$ and dash-dotted lines the mode $k = 1\,\mathrm{Mpc}^{-1}$.
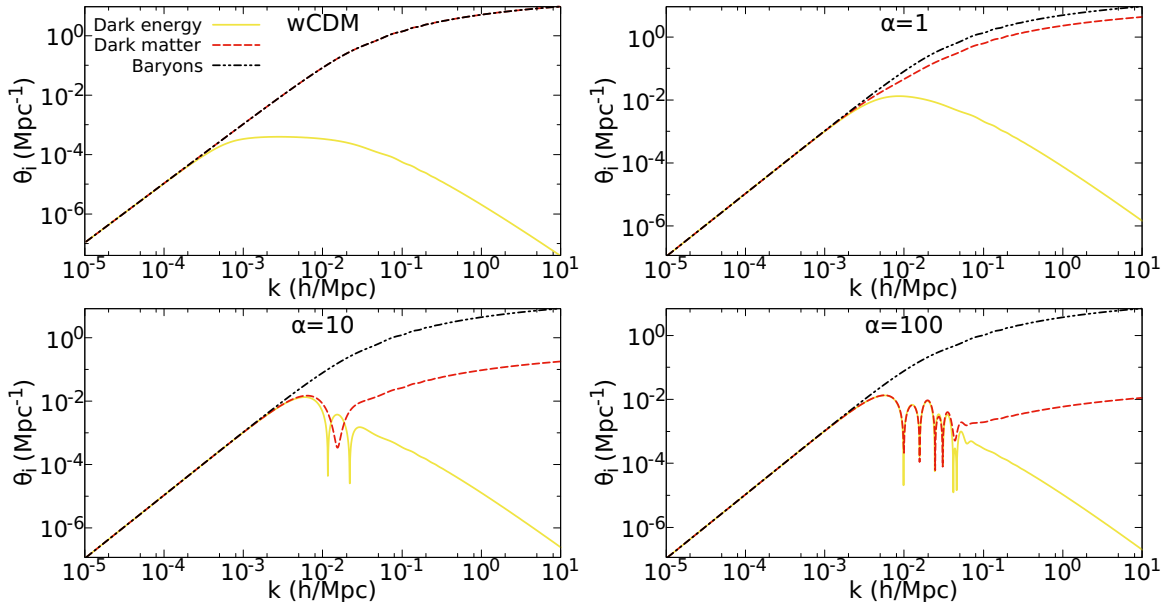

Figure 8: In these plots, we show the k-dependence of today's velocity of DE (yellow line) and DM (red line) for several values of the coupling constant α. As a reference for DM, we also plot baryons velocity (black line). We see that large scale behaviour is unaffected, since the interacting term vanishes. On intermediate scales DAO appear, showing the competition between the DE drag and the gravitational pull, resulting in a suppression of DM velocity. Finally, at small scales, DM and DE velocities decouple as consequence of the gravitational pull that dominates over the DE drag. These plots are done using the Newtonian gauge, where matter velocities dominate over DE ones.

simply introducing a cut-off scale so that the coupling becomes k-dependent as follows [2]

$$\alpha(k) = \alpha\, e^{-k/k_s}\ . \tag{3.1}$$

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[2] Although we have in mind existing screening mechanisms for certain types of theories like e.g. chameleon, symmetron, K-mouflage or Vainshtein (see for instance [45]), we do not intend our parameterisation to fully capture or provide a precise characterisation of those screening mechanisms. Rather, we want to explore if cutting-off small scales could improve the compatibility with data.
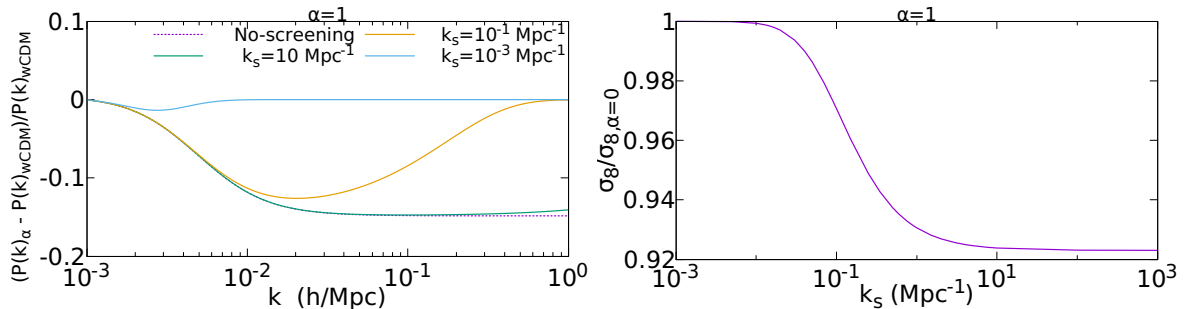
Figure 9: In the left plot, we show the relative ratio of the matter power spectrum using $\alpha = 1$ for several values of the screening scale k_s, including the previous model with no screening (dotted purple line). In the right plot, we display the ratio of the value of σ_8 for $\alpha = 1$ and the non interacting value $\sigma_{8,\alpha=0}$, as a function of the parameter k_s.

Focusing on the matter power spectrum and the parameter σ_8 to illustrate the effect of the cut-off we show the results in Figure 9. In the left panel, we show the relative ratio of the matter power spectrum fixing the parameter $\alpha = 1$ for several values of the cut-off scale k_s. As expected, we see that those modes smaller than the cut-off scale follow the purple dotted line that represents the previous model of Section 3.1, where no screening takes place. For those modes larger than k_s the matter power spectrum tends to behave as a non interacting wCDM model, while for scales around the cut-off there is an interpolation between both behaviours. In the case of the parameter σ_8, as it is the root mean square of fluctuation amplitude within spheres of scales $8\,h^{-1}$Mpc, we infer from the right plot of Figure 9 that when the coupling cut-off scales are much smaller than the σ_8-spheres scales its value is not affected by the cut-off and remains as in Section 3.1. When the cut-off scale includes inside it the σ_8-spheres no lowering on the value of σ_8 appears, as expected due to the scale-suppression of the coupling. It is important to remark that very large values of k_s would belong to the non-linear regime where non-linear effects are expected to take on. As before for those cut-off scales represented by k_s close to $8\,h^{-1}$Mpc we have an interpolation between both behaviours.

4 Fit to the data

Once we have explored the evolution of the cosmological perturbations and their effects on the CMB and matter power spectra for the considered interacting model, we turn our attention to observations in order to confirm and update the results obtained in [31]. We will also explore the relevance of introducing the aforementioned cut-off to discern if a given screening scale is preferred by data.

We use the public code of Markov chains Monte Carlo called MontePython [46, 47] applied to our modified `CLASS` code, to fit several cosmological parameters and our coupling parameters α and k_s with the available data. We use the full Planck 2018 likelihood [36, 40] containing data of high-ℓ and low-ℓ from CMB temperature (TT), polarisation (EE), cross correlation of temperature and polarisation (TE) and the CMB lensing power spectrum, the JLA likelihood with supernovae data [48], the BAO combined data [49–51], the PlanckSZ likelihood of Planck Sunyaev-Zeldovich effect [52] and the weak lensing likelihood CFHTLens [53]. The last two datasets are implemented as priors on $S_8^{\rm SZ} = \sigma_8(\Omega_{\rm m}/0.27)^{0.3}$

and $S_8^{\text{CFHTL}} = \sigma_8(\Omega_{\text{m}}/0.27)^{0.46}$ respectively. The use of PlanckSZ data presents some caveats that we will discuss below.

For our fit, we consider as cosmological parameters the baryon density defined as $100\ \Omega_{\text{b}}h^2$, the DM density as $\Omega_{\text{dm}}h^2$, the scalar spectral index n_s, the primordial amplitude as $10^9 A_s$, the reionisation optical depth τ_{reio}, the equation of state of DE as w, the angular acoustic scale as $100\ \theta_s$ and the coupling parameters for each model, α for the constant coupling model while α and k_s for the screening model. Also, we consider as derived parameters the redshift of reionisation z_{reio}, the Hubble parameter H_0, the matter fluctuation amplitude at $8h^{-1}\text{Mpc}$ as σ_8 and the total matter density Ω_{m}. The DE equation of state is assumed to be constant and restricted to $w > -1$ in order to avoid instabilities of the interacting model. For the same reason, we assume the coupling parameter α to be positive or, at least, with not a large value if negative (see Appendix B).

Finally, we remind that any significant difference found with the elastic coupling models will originate from the perturbation sector since we do not change the background cosmology.

4.1 Constant coupling model $\alpha = const$

Flat prior on $\log_{10} \alpha$

First, we will start our analysis by considering a flat prior on $\log_{10} \alpha$ over the range $\log_{10} \alpha \in [-8, 4]$. The purpose of this initial fit is to obtain a sense of the order of magnitude of α so we can explore later in more detail around the order of magnitude obtained here. One can argue we are excluding the non-interacting case $\alpha = 0$ but further analysis will support our choice. Furthermore, we know that values of α smaller than 10^{-2} do not give any appreciable deviation with respect to the non-interacting case, so this first analysis should allow us to infer, as a first approach, if the effects of the coupling are preferred by data or not.

In Table 1, we show the mean values with 1σ and 2σ confidence limits for all the parameters considered for the reference wCDM model and for the elastic coupling model when the logarithmic prior is used on α. In Figure 10, we display the one-dimensional posterior distribution and the two-dimensional contours for some parameters. Although there are significant differences in some cosmological and derived parameters, here we only focus on the coupling parameter α as we will study them in detail below when using a linear flat prior on α. With the result of the coupling parameter $\log_{10} \alpha = 0.003^{+0.15}_{-0.13}$, we confirm the ideas explained at the end of Section 2, where the normalisation of the coupling parameter was chosen to get a natural value $\alpha \sim \mathcal{O}(1)$. In view of this result we will perform the analysis with a flat prior on α incorporating the gained knowledge that $\alpha \sim \mathcal{O}(1)$ seems to be the preferred value by data. Furthermore, this analysis will permit to fully include the non interacting value of the coupling parameter $\alpha = 0$. We will show the consistency of the results obtained with both priors below, so we defer the discussion to the next section.

Flat prior on α

Once we have confirmed, using the logarithmic prior for α, that data seems to favour the coupling parameter to be $\alpha \sim \mathcal{O}(1)$, we will perform the same analysis but now using a flat prior on α instead of $\log_{10} \alpha$ over the range $\alpha \in [-0.1, 100]$, so we are now including the non interacting case. The range has been chosen sufficiently wide as to avoid any border effect in the sampling of the likelihood.

In Table 2, we show the constraints obtained for the reference model wCDM and for the interacting model for all cosmological, derived and model parameters. In Figure 11, we display for some parameters the one-dimensional posterior distribution and the two dimensional

	wCDM model			Elastic Interaction $\log_{10} \alpha$		
Param.	mean$\pm\sigma$	2σ lower	2σ upper	mean$\pm\sigma$	2σ lower	2σ upper
$100\Omega_{\mathrm{b}}h^2$	$2.264^{+0.015}_{-0.015}$	2.235	2.294	$2.243^{+0.016}_{-0.016}$	2.211	2.274
$\Omega_{\mathrm{dm}}h^2$	$0.1163^{+0.001}_{-0.001}$	0.1143	0.1183	$0.1193^{+0.0013}_{-0.0012}$	0.1169	0.1219
n_s	$0.9721^{+0.0042}_{-0.0043}$	0.9639	0.9807	$0.9662^{+0.0043}_{-0.0045}$	0.9576	0.9753
$10^9 A_s$	$2.063^{+0.035}_{-0.032}$	1.993	2.133	$2.105^{+0.034}_{-0.035}$	2.038	2.177
τ_{reio}	$0.0502^{+0.0092}_{-0.0082}$	0.03266	0.0686	$0.05618^{+0.0079}_{-0.0087}$	0.03955	0.07311
w	$-0.9478^{+0.022}_{-0.043}$	-0.9999	-0.8879	$-0.9788^{+0.006}_{-0.021}$	-0.9999	-0.9389
$100\,\theta_s$	$1.042^{+0.0003}_{-0.0003}$	1.042	1.043	$1.042^{+0.00032}_{-0.00032}$	1.041	1.043
z_{reio}	$7.141^{+0.89}_{-0.86}$	5.298	8.93	$7.837^{+0.82}_{-0.82}$	6.202	9.491
$H_0\ [\frac{\mathrm{km}}{\mathrm{s\,Mpc}}]$	$67.88^{+1.2}_{-0.96}$	65.71	69.97	$67.58^{+0.85}_{-0.65}$	66.02	69.15
σ_8	$0.7898^{+0.012}_{-0.009}$	0.7693	0.8094	$0.7516^{+0.012}_{-0.012}$	0.7272	0.7764
Ω_{m}	$0.3018^{+0.009}_{-0.012}$	0.2817	0.3226	$0.3104^{+0.008}_{-0.010}$	0.2933	0.3287
$\log_{10} \alpha$	$-$	$-$	$-$	$0.003^{+0.15}_{-0.13}$	-0.283	0.283

Table 1: In this table, we show the mean and 1σ values and the 2σ limits for the cosmological and derived parameters for a wCDM model (left) and for the interacting model using a logarithmic analysis on the coupling parameter α (right).

contours obtained with the previously explained datasets. We find a significant increase in the DM density $\Omega_{\mathrm{dm}}h^2$ while a slight decrease in the baryons density $100\Omega_{\mathrm{b}}h^2$ appears, which combined they are expressed on a small increase of the total matter density Ω_{m}. We also find a small decrease of the value of the scalar spectral index n_s and an increase on the value of the primordial amplitude $10^9 A_s$. However, what it is really significant is the lowering in the σ_8 value if we compare the wCDM model and the elastic coupling model. This result can be connected to what we found in the previous sections: the interaction induces a coupling between DE and DM that appears at the late universe. This coupling leads to the freeze of the growth of structures represented by the density contrast of DM in Figure 4 and, as a consequence, we have the suppression at small scales, where the coupling is efficient, as shown in Figure 1. The final consequence is a lower value of σ_8. Even more remarkable than the previous result is the fact that a non-zero value of the coupling strength is obtained at more than 3σ confidence level as $\alpha \underset{3\sigma}{\in} [0.249, 1.899]$. These results confirm the ones obtained in Ref. [31] but now with updated Planck 2018 data and including polarisation, and they
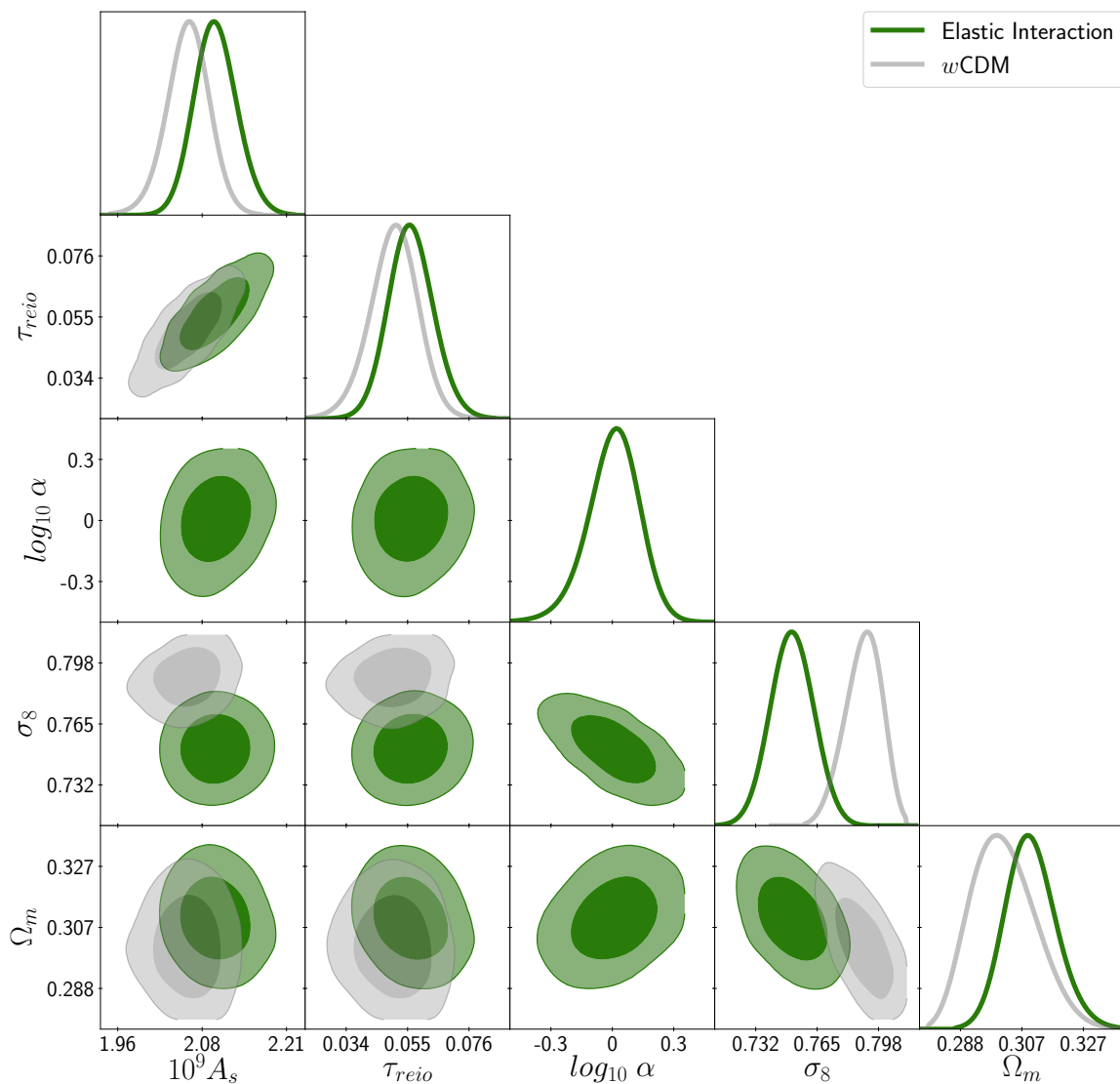
Figure 10: The one-dimensional 1σ posteriors distributions and the 2σ contours of several parameters for the constant coupling model using a logarithmic analysis on the model parameter (green) and the reference wCDM model (grey).

are also similar to the ones in Ref. [43], where the same coupling is used between DE and baryons. The results obtained here are compatible with those we got using the logarithmic parameter $\log_{10}\alpha$ as one would expect.

Finally, we find a significant $\Delta\chi^2 = 23.4$ improvement with respect to the wCDM model. This improvement is obtained by comparing the best fit of the interacting model with the best fit of a non-interacting wCDM[3]. However, this is not enough to claim that the model is preferred by data with a strong statistical significance, since adding more parameters will

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[3]It may be worth noticing that a pure ΛCDM will give a $\Delta\chi^2$ at least as big as wCDM so our discussion here is also valid for the pure ΛCDM case. As a matter of fact, we have explicitly compared to the pure ΛCDM and we have found $\Delta\chi^2 = 24.1$. In this respect, although ΛCDM has one fewer parameter, this is not sufficient to compensate for the big χ^2 difference.

Param.	wCDM model			Elastic Interaction α		
	mean$\pm\sigma$	2σ lower	2σ upper	mean$\pm\sigma$	2σ lower	2σ upper
$100\Omega_{\rm b}h^2$	$2.264^{+0.015}_{-0.015}$	2.235	2.294	$2.241^{+0.015}_{-0.016}$	2.211	2.272
$\Omega_{\rm dm}h^2$	$0.1163^{+0.001}_{-0.001}$	0.1143	0.1183	$0.1194^{+0.0012}_{-0.0012}$	0.1171	0.1218
n_s	$0.9721^{+0.0042}_{-0.0043}$	0.9639	0.9807	$0.9659^{+0.0042}_{-0.0045}$	0.9575	0.9748
$10^9 A_s$	$2.063^{+0.035}_{-0.032}$	1.993	2.133	$2.115^{+0.032}_{-0.034}$	2.048	2.183
$\tau_{\rm reio}$	$0.0502^{+0.0092}_{-0.0082}$	0.03266	0.0686	$0.05817^{+0.0081}_{-0.0082}$	0.04153	0.07468
w	$-0.9478^{+0.022}_{-0.043}$	-0.9999	-0.8879	$-0.9814^{+0.004}_{-0.019}$	-0.9999	-0.9426
$100\,\theta_s$	$1.042^{+0.0003}_{-0.0003}$	1.042	1.043	$1.042^{+0.00032}_{-0.00032}$	1.041	1.043
$z_{\rm reio}$	$7.141^{+0.89}_{-0.86}$	5.298	8.93	$8.051^{+0.8}_{-0.78}$	6.459	9.694
$H_0\ [\frac{\rm km}{\rm sMpc}]$	$67.88^{+1.2}_{-0.96}$	65.71	69.97	$67.1^{+0.77}_{-0.62}$	65.63	68.46
σ_8	$0.7898^{+0.012}_{-0.009}$	0.7693	0.8094	$0.7455^{+0.012}_{-0.012}$	0.7215	0.7688
$\Omega_{\rm m}$	$0.3018^{+0.009}_{-0.012}$	0.2817	0.3226	$0.3166^{+0.0075}_{-0.0087}$	0.301	0.3338
α	$-$	$-$	$-$	$1.005^{+0.26}_{-0.33}$	0.4374	1.648

Table 2: In this table we show the constraints obtained from the data explained in the main text for the wCDM model (left) and for the interacting model with the coupling parameter constant (right).

always lead to an improvement in the χ^2-squared value, therefore we will resort to a more informative criterion. In particular, we will apply the AIC criterion [54] based on the value of the quantity defined as

$$AIC \equiv -2\log\mathcal{L} + 2k\,, \tag{4.1}$$

where $\log\mathcal{L}$ is the maximum of the likelihood and k the number of parameters used in the fit. Applying this criterion, we find an improvement of $\Delta AIC = 21.4$, which in turn constitutes a strong evidence favouring the interacting model. This confirms that the substantially lower value of the best fit χ^2 for the interacting model does indicate and the price of adding only one new parameter (two in the case of a pure ΛCDM) is not a sufficient penalty.

Analysis for each data-set

Due to the more than 3σ preference and the enormous improvement in the χ^2 for the elastic coupling model, we will analyse the impact of each dataset on the determination of the cosmological parameters and the model parameter α as well as their contribution to lowering
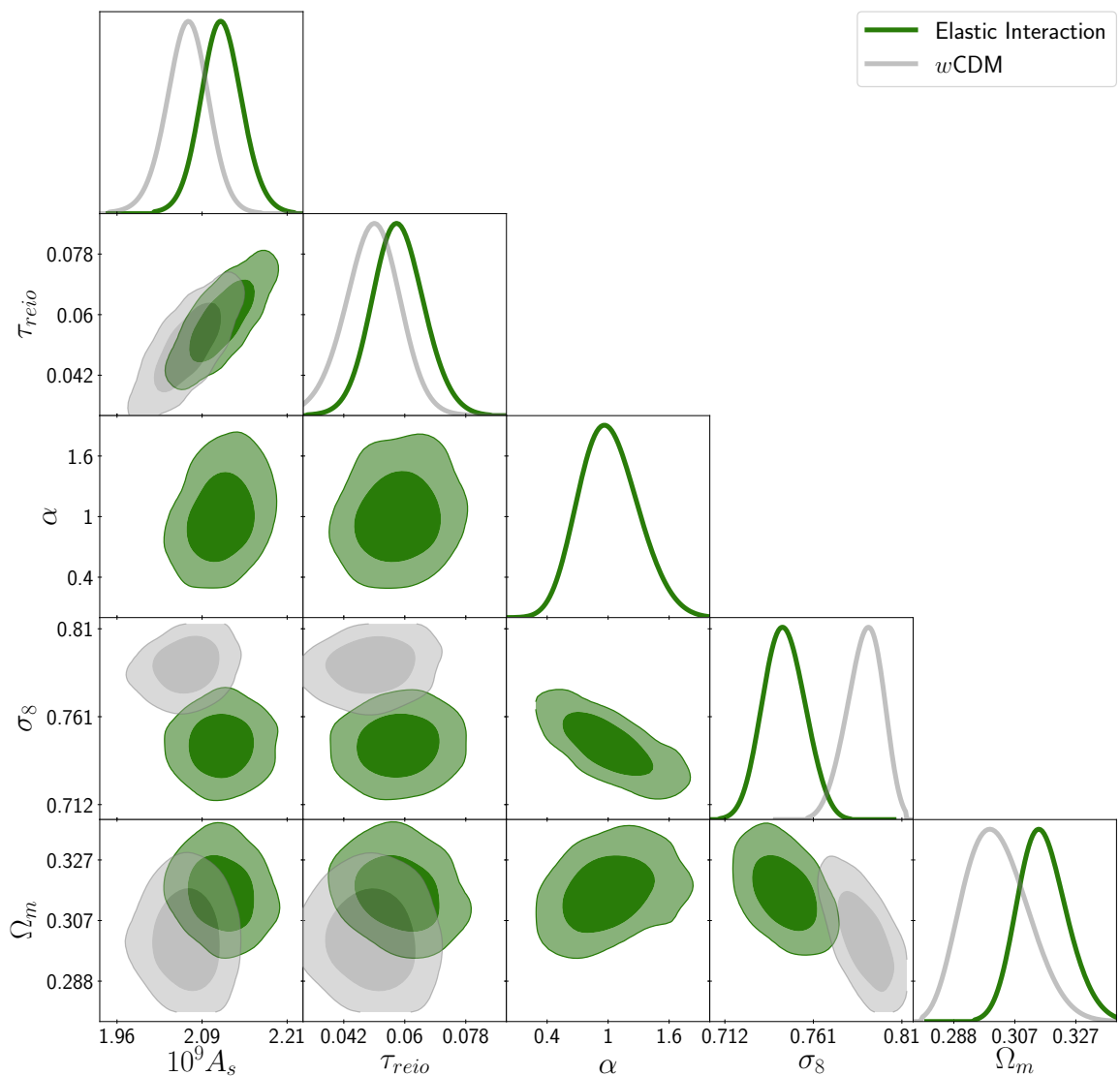
Figure 11: The two-dimensional 1σ and 2σ posterior of several parameters for the constant coupling model.

the best-fit χ^2. In this analysis we use the combinations of Planck 2018 likelihood of TT, EE, TE and lensing (P18), Planck 2018 with JLA and BAO likelihood (P18+JLA+BAO), Planck 2018 plus JLA and BAO with the CFHTLenS likelihood (P18+JLA+BAO+CFHTL) and Planck 2018 plus JLA and BAO with the Sunyaev-Zeldovich Planck likelihood (P18+JLA+ BAO+SZ). We find no substantial change on the previous parameters with the exception of the elastic coupling parameter α and the derived parameter σ_8, which are shown in Table 3 and Figure 13 . From this analysis, we infer that the Sunyaev-Zeldovich Planck likelihood is the main contributor to the possible detection of the elastic coupling when combined with Planck18, BAO and JLA data, with some additional contribution from the CFHTLenS likelihood. The Planck 2018 data alone, or with the JLA and BAO data, cannot establish a lower limit to the coupling parameter which would suggest a detection, but they are only able to put an upper limit that is compatible (at least at the 2σ level as one can see from Figure 13)
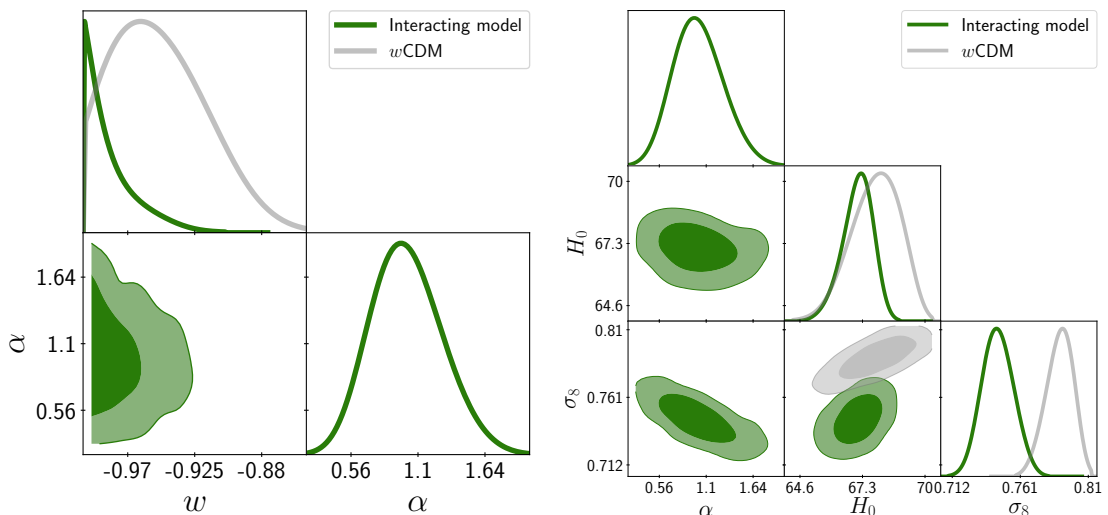
Figure 12: The two-dimensional 1σ and 2σ posteriors for the $w - \alpha$ plane (left) and for the model parameter α, the Hubble constant H_0 and σ_8 (right).

with the combined data or with P18+JLA+BAO+SZ combined analysis. For completeness, we also show the result for the combination of parameters $\sigma_8 \left(\frac{\Omega_m}{0.27}\right)^{0.3}$. It shows that when the interaction is not preferred by the data set, $\alpha \to 0$, both σ_8 and $\sigma_8 \left(\frac{\Omega_m}{0.27}\right)^{0.3}$ tend to have the standard model value, while when the value of the coupling parameter α shows the interaction is preferred, $\alpha \to 1$, both σ_8 and $\sigma_8 \left(\frac{\Omega_m}{0.27}\right)^{0.3}$ lower their values, although they are compatible at least at the 2σ level as shown in Figure 13. This result is in agreement with the intuitive expectation that the coupling suppresses the growth of structures so the interaction indeed prefers lower values of σ_8 (for constant Ω) that are in better agreement with low-redshift data.

As we have seen, the detection of the interaction parameter is driven by PlanckSZ data. There are some caveats, however, as to the consistency of using these data as a prior on $\sigma_8 \left(\frac{\Omega_m}{0.27}\right)^{0.3}$. To obtain this value, a translation from temperature profiles to density is necessary as well as a proper identification of the mass bias or mass calibration (see e.g. [55] and references therein) that require the use of hydrodynamical simulations properly tackling the non-linear scales and this is done by using ΛCDM-based simulations. We lack any devoted analysis of the non-linear effects for the model under consideration so it might not be justified to use those data. Our results are thus only valid as long as the potential effects from the interaction are sufficiently small on those scales or, at least, below the precision of the derived values. In any case, it is intriguing that PlanckSZ data as used here not only substantially improves the fit with respect to wCDM, but it crucially drives the potential detection of α so our findings are worth being reported here.

The obtained results motivate to consider two fiducial cosmologies in the forecast analysis that we will perform in Sec. 5, one with a non interacting value of the coupling parameter $\alpha = 0$ and with the combined result obtained for it $\alpha = 1.005^{+0.26}_{-0.33}$. Using both fiducial cosmologies will allow us to avoid possible bias effects in our forecast since the potential detection of $\alpha \neq 0$ is clearly driven by the PlanckSZ data.

	P18	P18+JLA+BAO	P18+JLA+BAO+CFHTL	P18+JLA+BAO+SZ	All data-sets
$\log_{10} \alpha$	$-4.2^{+1.7}_{-3.5}$	$-4.1^{+2.2}_{-2.9}$	$-3.1^{+3.8}_{-1.8}$	$-0.01^{+0.15}_{-0.12}$	$0.003^{+0.15}_{-0.13}$
σ_8	$0.7965^{+0.028}_{-0.019}$	$0.8123^{+0.013}_{-0.02}$	$0.798^{+0.03}_{-0.014}$	$0.7527^{+0.012}_{-0.012}$	$0.7516^{+0.012}_{-0.012}$
$\sigma_8(\frac{\Omega_m}{0.27})^{0.3}$	$0.8497^{+0.017}_{-0.029}$	$0.8485^{+0.04}_{-0.013}$	$0.8313^{+0.039}_{-0.021}$	$0.7856^{+0.012}_{-0.012}$	$0.7836^{+0.012}_{-0.011}$

Table 3: In this table, we show constraints on α, σ_8 and the parameter $\sigma_8(\frac{\Omega_m}{0.27})^{0.3}$ for the elastic model obtained from the different combination of data-sets: Planck 2018 likelihood of TT, EE, TE and lensing data (P18), Planck 2018 with JLA and BAO likelihood (P18+JLA+BAO), Planck 2018 plus JLA and BAO with the CFHTLenS likelihood (P18+JLA+BAO+CFHTL), Planck 2018 plus JLA and BAO with the Sunyaev-Zeldovich Planck likelihood (P18+JLA+BAO+SZ) and all the data-sets. The other cosmological parameters do not suffer from any substantial change.
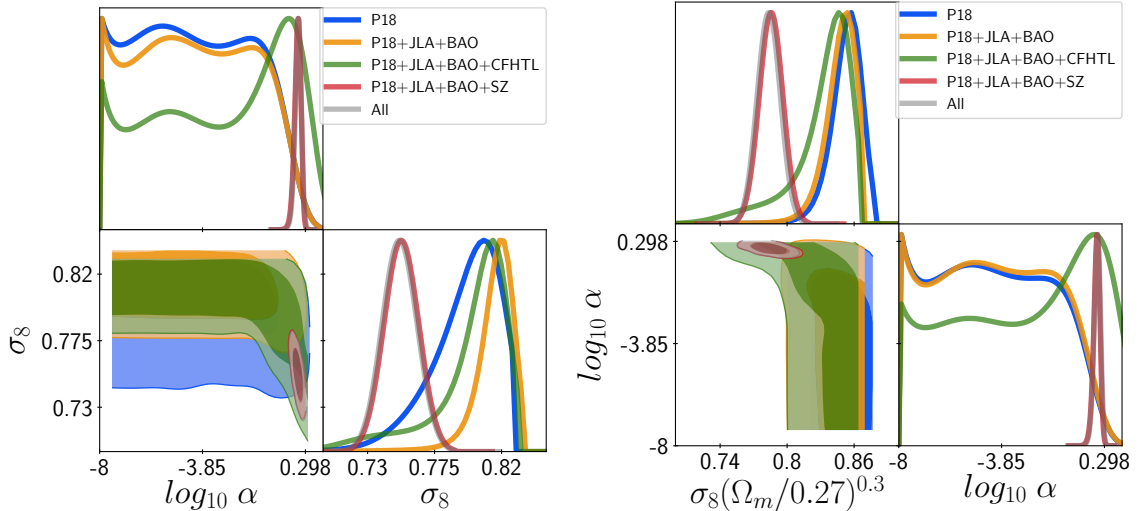
Figure 13: The two-dimensional 1σ and 2σ posterior of the coupling parameter α and the parameter σ_8 (left) and for the coupling parameter α and the parameter $\sigma_8(\frac{\Omega_m}{0.27})^{0.3}$ (right) for the constant coupling model, using different combinations of data-sets: P18 means Planck 2018 likelihood of TT, EE, TE and lensing data (blue), P18+JLA+BAO means Planck 2018 with JLA and BAO likelihood (yellow), P18+JLA+BAO+CFHTL denotes Planck 2018 plus JLA, BAO and CFHTLenS data (green), P18+JLA+BAO+SZ refers to Planck 2018 plus JLA, BAO and Sunyaev-Zeldovich Planck likelihood (red) while All uses all the previous data-sets (grey).

4.2   Screening model $\alpha e^{-k/k_s}$

In this second scenario, we will assume that the coupling strength has the scale dependence explained in Sec. 3.2 which allows to have the UV cut-off scale for the interaction as $\alpha(k) = \alpha e^{-k/k_s}$. We will employ a similar methodology as in the previous Section, i.e., the

normalisation for the coupling chosen so that $\alpha \sim \mathcal{O}(1)$ and, consequently, we will set a flat prior over $\alpha \in [-0.1, 100]$, while for the scale cut-off parameter we will use a flat prior on $\log_{10} k_s \in [-4, 21]$.

In Table 4, we show the best fit value, the mean with 1σ confidence limits and also 2σ upper and lower limits for the cosmological and model parameters, both for the reference model wCDM and the screening model. The cosmological parameters constraints are essentially indistinguishable from the ones obtained for the constant coupling model. The coupling parameter α presents a slightly higher value but again is more than 3σ away from a non interacting value. The screening scale k_s is very poorly constrained but clearly prefers very small scales. This points towards no scale suppression preferred by present data. In this respect, we should mention that those small values of k_s are nevertheless beyond the regime of applicability of our analysis. Not only would we expect a different behaviour of the screening on those scales, but non-linear effects will also become important and they should be properly incorporated. Moreover, those very small scales belong to regions not explored by the used data. Finally, we found a significant $\Delta\chi^2 = 27.1$ improvement with respect to the wCDM model. As we again find the coupling is strongly preferred with a huge improvement in the value of χ^2 and not a scale cut-off, at least for cosmological scales, we conclude that the elastic coupling model is preferred by the current data, following the results of [31] with improved data.

5 Galaxy surveys forecast

The results obtained in the preceding Section show how the scenario with the elastic coupling is statistically preferred by data and potentially rule out the standard model at more than 3σ for some combinations of cosmological data that include low-redshift measurements of clustering. In the upcoming years, several new galaxy surveys are planned that will further improve on their predecessors, being able to constrain the cosmological parameters with even higher precision. In view of the crucial impact that those measurements might have on the determination of the interaction, we will next proceed to estimate the discriminating power of these new experiments for the elastic coupling scenarios under consideration in this work. We will focus on the Javalambre Physics of the Accelerating universe Astrophysical Survey (J-PAS) [8], but we will also consider DESI [3] and Euclid [2] for comparison and to estimate the forecast performance of different planned surveys. This forecast analysis will be done using the publicly available code `FARO` [56], which is a galaxy survey forecast code to compute Fisher matrices for galaxy clustering and weak lensing observables. As input for `FARO` we will need the derivatives of the corresponding observables (to be defined below) with respect to the parameters that we want to forecast (α in our case) evaluated on the fiducial model. These derivatives will be obtained from the numerical outputs of the modified `CLASS` code that includes the interaction. Before entering into the Fisher analysis employed for our forecast, let us briefly introduce the main survey of interest for us, i.e., J-PAS.

5.1 The J-PAS survey

J-PAS is a spectro-photometric survey that will be conducted at the Observatorio Astrofísico de Javalambre (OAJ) by the 2.5m diameter Javalambre Survey Telescope (JST/ T250) equipped with the 5 sq. deg. Field of View Javalambre Panoramic Camera (JP-Cam). The filter system includes 54 narrow- and 4 broad-band filters covering the optical range [57] which will allow J-PAS to measure the redshift of several millions of Luminous Red Galaxies

	wCDM model			Elastic Interaction $\alpha e^{-k/k_s}$		
Param.	mean$\pm\sigma$	2σ lower	2σ upper	mean$\pm\sigma$	2σ lower	2σ upper
$100\Omega_b h^2$	$2.264^{+0.015}_{-0.015}$	2.235	2.294	$2.240^{+0.015}_{-0.016}$	2.209	2.27
$\Omega_{\rm dm} h^2$	$0.1163^{+0.0010}_{-0.0010}$	0.1143	0.1183	$0.1196^{+0.0012}_{-0.0011}$	0.1172	0.1219
n_s	$0.9721^{+0.0042}_{-0.0043}$	0.9639	0.9807	$0.9656^{+0.0040}_{-0.0045}$	0.9574	0.9747
$10^9 A_s$	$2.063^{+0.035}_{-0.032}$	1.993	2.133	$2.101^{+0.031}_{-0.036}$	2.036	2.169
$\tau_{\rm reio}$	$0.0502^{+0.0092}_{-0.0082}$	0.0327	0.0686	$0.0552^{+0.0076}_{-0.0085}$	0.0392	0.0722
w	$-0.948^{+0.022}_{-0.043}$	-0.9999	-0.8879	$-0.979^{+0.004}_{-0.021}$	-0.9999	-0.9427
$100\,\theta_s$	$1.042^{+0.00030}_{-0.00031}$	1.042	1.043	$1.042^{+0.00031}_{-0.00031}$	1.041	1.043
$z_{\rm reio}$	$7.14^{+0.89}_{-0.86}$	5.298	8.93	$7.75^{+0.76}_{-0.84}$	6.131	9.376
$H_0\ [\frac{\rm km}{\rm s\,Mpc}]$	$67.88^{+1.20}_{-0.96}$	65.71	69.97	$67.44^{+0.73}_{-0.65}$	65.98	68.92
σ_8	$0.7898^{+0.012}_{-0.0093}$	0.7693	0.8094	$0.7483^{+0.011}_{-0.012}$	0.7253	0.7723
$\Omega_{\rm m}$	$0.3018^{+0.009}_{-0.012}$	0.2817	0.3226	$0.3123^{+0.008}_{-0.009}$	0.2953	0.3289
α	$-$	$-$	$-$	$1.13^{+0.26}_{-0.33}$	0.53	1.78
$\log_{10} k_s [{\rm Mpc}^{-1}]$	$-$	$-$	$-$	$8.5^{+5.2}_{-5.0}$	-1.1	16.6

Table 4: In this table we give the obtained constraints for the cosmological and derived parameters for a wCDM model (left) and for the interacting model with the coupling parameter constant screened (right).

(LRG), Emission Line Galaxies (ELG) up to $z = 1.3$ and quasars (QSO) up to $z = 3.9$ with a precision of $\delta z = 0.003(1 + z)$. The survey will cover a maximum area of 8500 sq. deg. in the northern hemisphere.

The optical and filter systems have already been tested in the mini-J-PAS survey which covered 1 sq. deg. and whose results have been recently published in [58]. These results have confirmed that the nominal photo-z accuracy is achieved and that the expected number density of galaxies allows for excellent cosmological measurements. Preliminary forecast analysis [18, 59, 60] showed that thanks to the large number of ELG and LRG detectable, J-PAS will perform very well in the redshift range $0.3 < z < 0.6$ and also will cover the high-redshift range thanks to the remarkable number of QSO.

J-PAS will have the unique capability of combining shear measurements of galaxy weak lensing, with (multitracer) clustering observations from redhsift-space distortions and Alcock-

Paczynski effects which will largely improve the precision of cosmological parameters measurements [56].

5.2 Fisher formalism for galaxy clustering

After this brief introduction of J-PAS, let us turn to explaining the methodology of our forecast analysis. Following the formalism explained in Ref. [60] (where we refer for more details), the galaxy power spectrum for multitracer (i and j) analysis can be defined as

$$P_{ij}(k_r, \hat{\mu}_r, z) = \frac{D_{A\,r}^2 E}{D_A^2 E_r} \left(A_i + R\hat{\mu}^2 \right) \left(A_j + R\hat{\mu}^2 \right) \hat{P}(k) e^{-k_r^2 \hat{\mu}_r^2 \sigma_r^2} , \qquad (5.1)$$

where the angular distance D_A and the Hubble function $E(z) \equiv H(z)/H_0$ build the prefactor due to Alcock-Paczynski effect [61], the parameters A_i and R are defined as $A_i = Db_i\sigma_8$ and $R = Df\sigma_8$ where D is the growth factor, f the growth rate, b_i the bias of each tracer and σ_8 is the amplitude of matter perturbations at $8\,h^{-1}\,\mathrm{Mpc}$. Also, $\hat{P}(k)$ is defined as $\hat{P}(k) = P(k)/\sigma_8^2$ and $\sigma_r = \delta z(1+z)/H(z)$, with δz the photometric redshift error of each survey. The variables are the redshift z, the scale k and the angle of separation with the line of sight $\hat{\mu}$. The subindex r means it is evaluated at the fiducial cosmology. Also because of the Alcock-Paczynski effect they are related as

$$k = Qk_r \quad \text{and} \quad \hat{\mu} = \frac{E}{E_r Q}\hat{\mu}_r, \qquad (5.2)$$

with

$$Q = \frac{\sqrt{D_A^2 \chi^2 \hat{\mu}_r^2 - D_{A\,r}^2 \chi_r^2(\hat{\mu}_r^2 - 1)}}{D_{A\,r}\chi}, \qquad (5.3)$$

being $\chi = \int_0^z \frac{dz}{H(z)}$ the comoving distance.

As shown in Ref. [60], the Fisher matrix for clustering for different tracers i and j centred at redshift bin z_a is

$$F_{\alpha\beta}^C(z_a) = \frac{V_a}{8\pi^2} \int_{-1}^1 \mathrm{d}\hat{\mu} \int_{k_{\min}}^\infty \mathrm{d}k k^2 \frac{\partial P_{ij}(k, \hat{\mu}, z_a)}{\partial p_\alpha}\bigg|_r C_{jl}^{-1} \frac{\partial P_{lm}(k, \hat{\mu}, z_a)}{\partial p_\beta}\bigg|_r C_{mi}^{-1} e^{-k^2 \Sigma_\perp^2 - k^2 \hat{\mu}^2 (\Sigma_\parallel^2 - \Sigma_\perp^2)}, \qquad (5.4)$$

where V_a is the total volume of the a-th bin, $k_{\min}$ is fixed to $0.007h\,\mathrm{Mpc}^{-1}$, $\{p_\alpha\}$ are the set of cosmological parameters we are doing the forecast for, C_{ij} is the data covariance matrix defined as $C_{ij} = P_{ij} + \frac{\delta_{ij}}{\bar{n}_i}$, with $\bar{n}_i$ the mean galaxy density of tracer i at z_a. Finally, the exponential suppression factor is introduced to remove the non-linear scales from the analysis, since properly including those scales necessitates an appropriate characterisation of the non-linear clustering that we lack, and the functions $\Sigma_\parallel$ and $\Sigma_\perp$ are defined as

$$\Sigma_\perp(z) = 0.785D(z)\Sigma_0 , \qquad (5.5)$$
$$\Sigma_\parallel(z) = 0.785D(z)\Big[1 + f(z)\Big]\Sigma_0 , \qquad (5.6)$$

where $\Sigma_0 = 11\,\mathrm{Mpc}/h$. We should bear in mind that this parameterisation was obtained for the standard ΛCDM model so, in principle, it is not directly applicable to our scenario with the elastic interaction because it gives a substantially different linear clustering evolution. Although it is expected that the non-linear clustering will also differ from ΛCDM, thus casting

doubts on the use of the above parameterisation, we do not expect the non-linear scale to differ significantly between both models. Hence, it is justified to use the above parameterisation also for the interacting model. Furthermore, the results of the Fisher analysis for our forecasts are not expected to be very sensitive to these details and it is then sensible to use the standard model values. For the real analysis with data, the difference can be more relevant and care should be taken in order to optimise the amount of information extracted from data.

In our analysis we use a binned distribution in scale k from $0.007h\,\mathrm{Mpc}^{-1}$ to $10h\,\mathrm{Mpc}^{-1}$ with 20 bins logarithmically equispaced, in the case of the angle of line of sight $\hat{\mu}$ from -1 to 1 with 200 bins equispaced, while the index i and j are the tracers: Luminous Red Galaxies (LRG), Emission Line Galaxies (ELG), Quasars (QSO) and Bright Galaxies (BGS). The redshift distribution depends on the experiment and are explained for the three surveys used for clustering J-PAS, DESI and spectroscopic Euclid in Appendix C.

For our fiducial cosmology we consider a flat universe with, $\Omega_{\mathrm{m}} = 0.310$, $h = 0.6774$, $w = -0.98$ and $n_s = 0.96$. Due to the more than 3σ result obtained in the previous sections we consider two fiducial cosmologies, one with $\alpha = 0$ and the other with $\alpha = 1$. In addition, as we will consider different tracers for the galaxy clustering power spectra, we need the fiducial values for biases. Since in the early universe dark matter decouples from baryons due to electromagnetic interactions, there is a bias between galaxies and dark matter halos. This bias evolves with the galaxy formation and evolution, and also depends on selection effects of different galaxy surveys. We will consider the following fiducial biases for each tracer [60],

$$b_{\mathrm{LRG}} = \frac{1.7}{D(z)} \, , \tag{5.7}$$

$$b_{\mathrm{ELG}} = \frac{0.84}{D(z)} \, , \tag{5.8}$$

$$b_{\mathrm{QSO}} = 0.53 + 0.289(1+z)^2 \, , \tag{5.9}$$

$$b_{\mathrm{BSG}} = \frac{1.34}{D(z)} \, , \tag{5.10}$$

for J-PAS and DESI, while for Euclid $b_{\mathrm{ELG}} = \sqrt{1+z}$ [2]. Since the interacting scenario that we are exploring in this work can potentially modify the bias, we will perform two analysis. Firstly, we will fix the bias to the previous fiducial values. Secondly, we will marginalise over the bias as free parameter in each redshift bin z_a, in order to take into account the possible correlations of our desired parameters with the bias.

The photometric redshift error δz for J-PAS is $\delta z = 0.003$ for LRG, ELG and QSO, for DESI is $\delta z = 0.0005$ for BGS, LRG and ELG and $\delta z = 0.001$ for QSO, while for EUCLID is $\delta z = 0.001$ for ELG.

5.2.1 Clustering results

In this Section we present the forecast analysis for Ω_{m} and interaction parameter α for the surveys J-PAS [8], DESI [3] and Euclid [2] using clustering information. If we consider the results shown in Section 4 to be reliable, then, our fiducial cosmology is the one with $\alpha = 1$. In Figure 14 , we provide the 1σ contour error for the fiducial cosmology $\alpha = 1$ that can be measured using the surveys J-PAS using tracers ELG+LRG+QSO or without QSO (black lines), DESI with LRG+ELG+QSO+BGS or without BGS (blue lines) and Euclid using ELG (yellow line).

With all the three considered surveys, we have an expected constraint of order $\sim 5\%$ on our model parameter α if we consider the fiducial cosmology to have $\alpha = 1$. These results
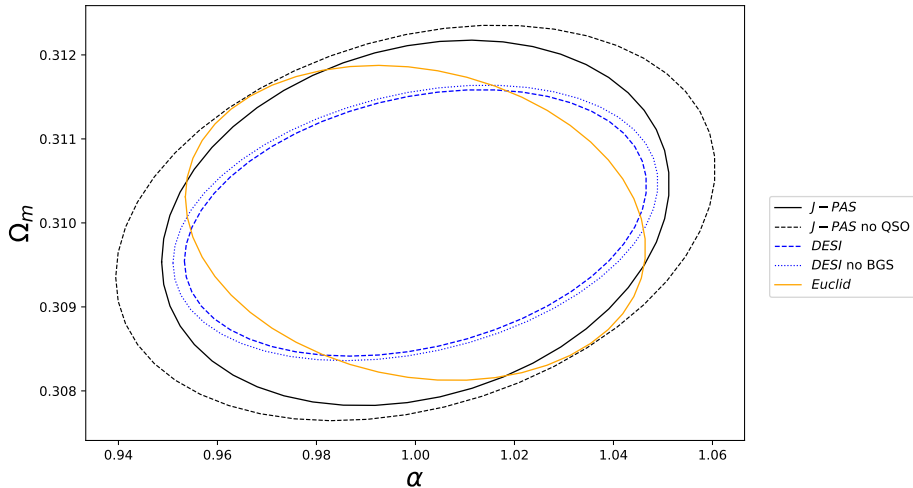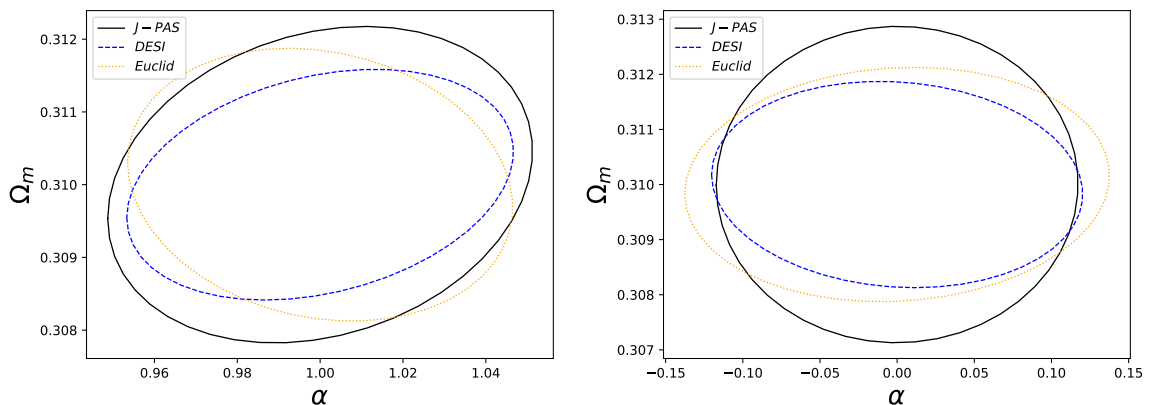
Figure 14: Using as fiducial cosmology $\alpha = 1$ and marginalising the bias, the 1σ contour errors for the parameters $\Omega_{\rm m}$ and α using clustering for the surveys J-PAS with ELG+LRG+QSO tracers (black solid line) and without QSO (black dashed line), for DESI with LRG+ELG+QSO+BGS (blue dashed line) and without BGS (blue dotted line) and for Euclid using ELG (yellow solid line).

would lead to a 10σ detection of the interaction and then, if the specifications of each survey are satisfied, such a detection is falsifiable with the data obtained in the next cosmological experiments. In particular, for the J-PAS survey, we can obtain the minimum area necessary to obtain a detection of $\alpha = 1$ at 5σ. Since the clustering Fisher matrix is proportional to the fraction of sky, we obtain that given a 1σ marginalised constraint σ_a for a parameter a for a given area $\deg_a^2$, we can calculate the new error for a new area as $\sigma_b = \sqrt{\deg_a^2 / \deg_b^2}\, \sigma_a$. Using this relation, we obtain that J-PAS needs only 983 $\deg^2$ to be able to provide a detection of $\alpha = 1$ at 5σ.

These results have been obtained marginalising the bias in each redshift bin to prevent effects from having a different clustering evolution in the presence of the interaction. Furthermore, we have fixed the fiducial cosmology to a non-standard model. Thus, in order to test the robustness of our analysis against these choices, we will now proceed to analyse the change in the expected errors against a change of the fiducial model and the bias marginalisation.

5.2.2 Forecast marginalising bias

In Figure 15, we show the expected 1σ contours for $\Omega_{\rm m}$ and α with a bias marginalisation in each redshift bin, for the fiducial cosmology both the standard one and the bestfit result obtained in Section 4. In Table 5 we show the marginalised errors for the surveys and fiducial cosmologies used. We find an improvement by a factor ~ 4 on the constraints for the interaction parameter α with respect to the results of Section 4. As we have marginalised the bias in each redshift, the obtained errors are higher than the ones we will obtain in the next section without marginalising. This is expected since the loss of information on the bias reduces the precision of the measurements.

Figure 15: Using the fiducial cosmology with $\alpha = 1$ (left) and with $\alpha = 0$ (right) and marginalising the bias, the 1σ contour errors for the parameters Ω_{m} and α using clustering data for the surveys J-PAS with ELG+LRG+QSO tracers (black solid line), DESI with LRG+ELG+QSO+BGS (blue dashed line) and Euclid using ELG (yellow dotted line).

Table 5: Marginalised errors of the parameters $\Delta\Omega_m$ and $\Delta\alpha$ using clustering data and marginalising the bias for J-PAS, Euclid and DESI surveys.

	Fiducial wCDM		Fiducial $\alpha = 1$	
Survey	$\Delta\Omega_m$	$\Delta\alpha$	$\Delta\Omega_m$	$\Delta\alpha$
J-PAS	0.0038	0.155	0.0029	0.068
Euclid	0.0028	0.182	0.0025	0.062
DESI	0.0025	0.159	0.0021	0.062

5.2.3 Forecast with standard bias

If we assume that the bias is not substantially modified by the interaction, it is justified to use the bias of the standard model introduced in Section 5.2. The benefit of not marginalising the bias is that we do not add error to our parameters Ω_{m} and α due to the loss of information. As before, in Figure 16 and in Table 6 we present the 1σ contour errors and the marginalised errors of the parameters Ω_m and α using as fiducial cosmology the bestfit one or the standard model for the surveys J-PAS, DESI and Euclid. Again, we find stronger constraints than using the currently available data used in Section 4, which can lead to a potential detection of this interaction.

5.3 Fisher formalism for weak lensing

As in the clustering forecast, we will follow the formalism explained in Ref. [60] for the weak lensing analysis. The relevant observable in this case is the convergence power spectrum for a redshift tomography analysis that is given by

$$P_{ij}(\ell) = H_0 \sum_a \frac{\Delta z_a}{E_a} K_i(z_a) K_j(z_a) \Omega_m^2 D(z_a)^2 \sigma_8^2 \hat{P}\left(\frac{\ell}{\chi(z_a)}\right) , \qquad (5.11)$$
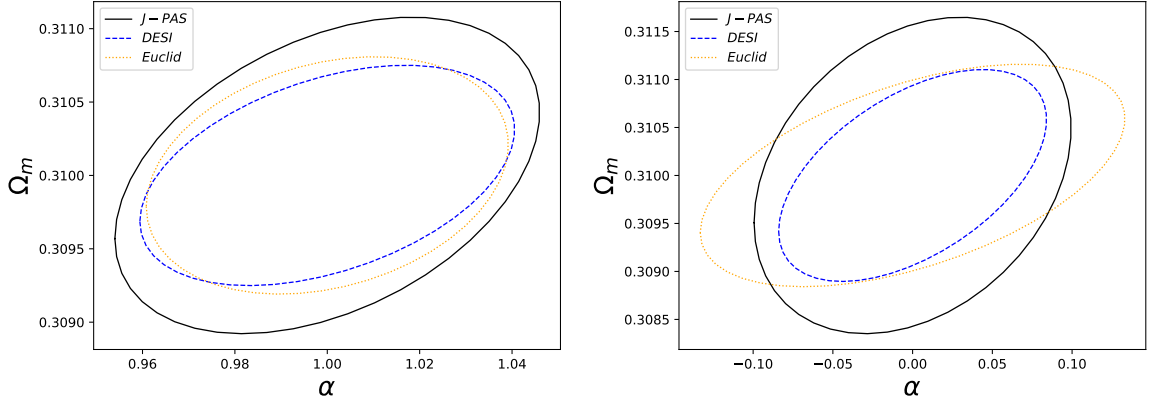
Figure 16: We show the 1σ contour errors for Ω_m and α for the fiducial cosmologies with $\alpha = 1$ (left) and with $\alpha = 0$ (right) and fixing the bias to the standard model one (5.10). These results correspond to using clustering data for the surveys J-PAS with ELG+LRG+QSO tracers (black solid line), DESI with LRG+ELG+QSO+BGS (blue dashed line) and Euclid using ELG (yellow dotted line).

Table 6: Marginalised errors of the parameters $\Delta\Omega_m$ and $\Delta\alpha$ using clustering data and bias fixed to the one in ΛCDM for J-PAS, Euclid and DESI surveys.

Survey	Fiducial wCDM		Fiducial $\alpha = 1$	
	$\Delta\Omega_m$	$\Delta\alpha$	$\Delta\Omega_m$	$\Delta\alpha$
J-PAS	0.0022	0.132	0.0014	0.061
Euclid	0.0015	0.176	0.0011	0.052
DESI	0.0015	0.111	0.0010	0.054

where $\hat{P}(k) = P(k)/\sigma_8^2$ and $D(z_a)$ is the growth factor, the index i and j refer to the bin in redshift and $\ell = k\chi(z)$. The functions K_i are defined as

$$K_i(z) = \frac{3H_0}{2}(1+z)\int_z^\infty \left(1 - \frac{\chi(z)}{\chi(z')}\right) n_i(z')dz' \,, \tag{5.12}$$

where $n_i(z)$ is the density function in the ith-bin of redshift such that considering photometric redshift error $\sigma_i = \delta z(1 + z_i)$

$$n_i(z) \propto \int_{\bar{z}_{i-1}}^{\bar{z}_i} z'^2 e^{-(z'/z_p)^{3/2}} e^{-\frac{(z'-z)^2}{2\sigma_i^2}}\, dz \,, \tag{5.13}$$

normalised for each n_i as $\int_0^\infty n_i(z)dz = 1$. z_p is defined as $z_p = z_{\text{mean}}/\sqrt{2}$, with z_{mean} the mean redshift of each survey. Finally, the Fisher matrix for weak lensing is

$$F_{\alpha\beta}^C = f_{\text{sky}} \sum_\ell \Delta\ln\ell \frac{(2\ell+1)\ell}{2} \text{Tr}\left[\frac{\partial\mathbf{P}}{\partial p_\alpha}\mathbf{C}^{-1}\frac{\partial\mathbf{P}}{\partial p_\beta}\mathbf{C}^{-1}\right] \,, \tag{5.14}$$
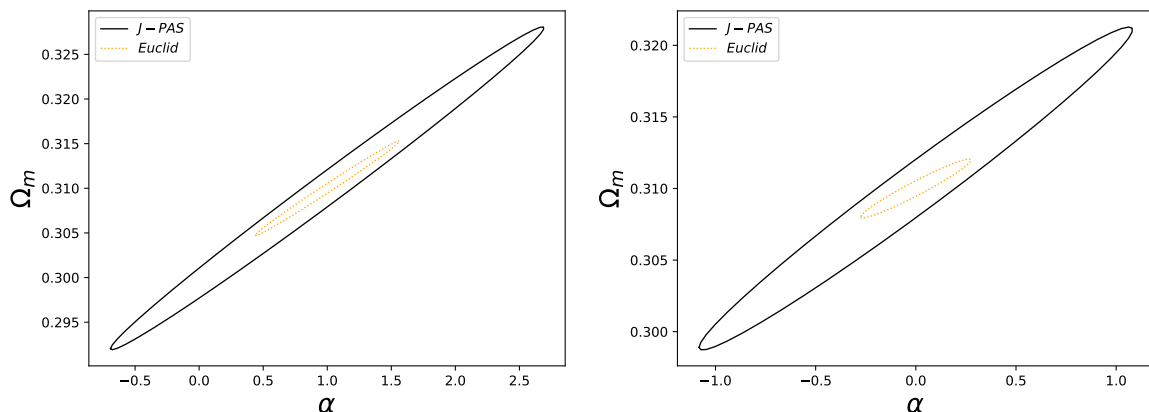
Figure 17: Using the fiducial cosmology with $\alpha = 1$ (left) and with $\alpha = 0$ (right), the 1σ contour errors for the parameters Ω_m and α using lensing data for the surveys J-PAS with ELG+LRG tracers (black solid line), and Euclid using ELG (yellow dotted line).

with $C_{ij} = P_{ij} + \gamma_{int}^2 \hat{n}_i^{-1} \delta_{ij}$, where the intrinsic ellipticity is $\gamma_{int} = 0.22$ and the number of galaxies per steradian in the ith-bin $\hat{n}_i$ is defined as

$$\hat{n}_i = n_\theta \frac{\int_{\bar{z}_{i-1}}^{\bar{z}_i} n_i(z) dz}{\int_0^\infty n_i(z) dz} \ , \tag{5.15}$$

where n_θ is the areal galaxy density, which for J-PAS is $n_\theta = 12.32$ and for Euclid $n_\theta = 30$.

The ith-bin distribution depends on the survey used for the lensing analysis (J-PAS and photometric Euclid) and they are explained in Appendix C, while for the scale distribution ℓ we use $\Delta \ln \ell = 0.1$ from $\ell_{\min} = 5$ and $\ell_{\max} = 906$. The mean redshift is $z_{\mathrm{mean}} = 0.5$ for J-PAS and $z_{\mathrm{mean}} = 0.9$ for Euclid, while the photometric redshift error is $\delta z = 0.03$ for J-PAS and $\delta z = 0.05$ for Euclid. As before, our two fiducial cosmologies are a flat universe with $\Omega_m = 0.310$, $h = 0.6774$, $w = -0.98$, $n_s = 0.96$ and both values of the coupling parameter $\alpha = 0$ and $\alpha = 1$ respectively.

5.3.1 Results weak lensing

In this section we present the expected constraints on the cosmological parameter Ω_m and on the model parameter α, for the surveys J-PAS [8] and Euclid [2] using lensing information. In Figure 17 we give the expected 1σ contour errors for both surveys while in Table 7 we show the marginalised errors, using as fiducial cosmology the bestfit result of Section 4 and the standard model. Although the constraints for Ω_m can be competitive, the constraints on α in the best scenario for Euclid and the standard cosmology are worse than the clustering ones or the results obtained in Section 4. J-PAS is even less capable of constraining the model parameter α.

The poor results obtained with lensing can be explained by the effect the model has in the growth function $D(z)$ and the lensing potential Ψ. For the values of α considered here, the impact on $D(z)$ and Ψ, in combination with the inability of this model to modify the slip parameter, are not enough to have competitive results for the errors with these surveys, specially for the case of J-PAS.

Table 7: Marginalised errors of the parameters $\Delta\Omega_m$ and $\Delta\alpha$ using lensing data for J-PAS and Euclid surveys.

	Fiducial wCDM		Fiducial $\alpha = 1$	
Survey	$\Delta\Omega_m$	$\Delta\alpha$	$\Delta\Omega_m$	$\Delta\alpha$
J-PAS	0.0149	1.43	0.0239	2.24
Euclid	0.0027	0.36	0.0070	0.74

6 Discussion and conclusions

In this work we have explored a cosmological scenario where dark matter and dark energy are coupled through their Euler equations, which arises from an interaction with no energy transfer. We have confirmed that these scenarios provide promising candidates for alleviating the existing tension in cosmological data when compared to the amplitude of the fluctuations inferred from CMB and those measured by low redshift galaxy surveys. It is remarkable that the latest cosmological data combined strongly favour the existence of an interaction with the non-interacting case excluded at more than 3σ. However, this detection is driven by PlanckSZ data, whose inclusion might not be well justified as we have explained. In any case, the obtained result clearly calls for a more detailed examination of these scenarios.

We have then proceeded to analyse the ability of future galaxy surveys to unveil the true presence of the interactions. Since one of the main effects of the interacting model is precisely a suppression of the clustering, these surveys might prove to be ideal for constraining these scenarios. Thus, we have used two observables based on clustering and weak lensing and we have compared the forecasts for J-PAS, DESI and Euclid. The analysis from clustering shows that the interaction will be definitely detected by these surveys and, despite their different specifications, all three of them will provide similar precision on the measurement of the interaction parameter, at the level of a few percent, with the non-interacting model clearly excluded. On the other hand, the weak lensing forecast provides a much poorer constraining power than clustering. The effects of this elastic model in the growth of structures that, at the end, modify the weak lensing are not enough to have competitive results with the surveys studied. Hence, clustering data will be the key for constraining the coupling parameter α.

Codes: Modified versions of the codes `CLASS` and `CAMB` for the computation of the evolution of linear perturbations are available on request.

A   Equations in synchronous gauge

The synchronous gauge is defined by the perturbed line element

$$ds^2 = a^2(\tau)\Big[- d\tau^2 + (\delta_{ij} + h_{ij}) dx^i dx^j \Big] , \tag{A.1}$$

where the perturbed spatial metric is written in terms of the scalar perturbations h and η as $h_{ij} = \mathrm{diag}(-2\eta, -2\eta, h + \eta)$. Following our previous notation, the conservation equations

governing the dark sector perturbations read

$$\delta'_{\rm dm} = -\left(\theta_{\rm dm} + \frac{1}{2}h'\right) , \tag{A.2}$$

$$\theta'_{\rm dm} = -\mathcal{H}\theta_{\rm dm} + \Gamma(\theta_{\rm de} - \theta_{\rm dm}) , \tag{A.3}$$

$$\delta'_{\rm de} = -3\mathcal{H}\left(c_{\rm s}^2 - w\right)\delta_{\rm de} - (1+w)\left(\theta_{\rm de} + \frac{1}{2}h'\right) + 9(1+w)\mathcal{H}^2\frac{c_{\rm s}^2 - w}{k^2}\theta_{\rm de} , \tag{A.4}$$

$$\theta'_{\rm de} = \left(-1 + 3c_{\rm s}^2\right)\mathcal{H}\theta_{\rm de} + \frac{k^2 c_{\rm s}}{1+w}\delta_{\rm de} - \Gamma R(\theta_{\rm de} - \theta_{\rm dm}) , \tag{A.5}$$

where a constant DE equation of state parameter is assumed and Γ and $R\Gamma$ conserve the same previous definitions.

B Brief discussion on stability

In this appendix we will give a brief description of the potential appearance of instabilities in the evolution of the perturbations due to the interaction. To that end, we first notice that the equations for the evolution of the perturbations can be combined to give a system of two second order equations involving only the density contrasts. If we consider sub-Hubble modes $k \gg \mathcal{H}$ during the matter dominated epoch, the resulting equations can be written as

$$\begin{bmatrix} \delta''_{\rm dm} \\ \delta''_{\rm de} \end{bmatrix} + \begin{bmatrix} \mathcal{H} + \Gamma & -\frac{\Gamma}{1+w} \\ -(1+w)\Gamma R & (1-3w)\mathcal{H} + R\Gamma \end{bmatrix} \begin{bmatrix} \delta'_{\rm dm} \\ \delta'_{\rm de} \end{bmatrix}$$
$$+ \begin{bmatrix} -\frac{3}{2}\mathcal{H}^2 & 3\frac{w-c_{\rm s}^2}{1+w}\mathcal{H}\Gamma \\ -\frac{3}{2}(1+w)\mathcal{H}^2 & c_{\rm s}^2 k^2 + 3(c_{\rm s}^2 - w)\mathcal{H}R\Gamma \end{bmatrix} \begin{bmatrix} \delta_{\rm dm} \\ \delta_{\rm de} \end{bmatrix} = 0 \tag{B.1}$$

where we have used that $R \gg 1$, $k^2\Phi \simeq -\frac{3}{2}\mathcal{H}^2\delta_{\rm dm}$ and $\mathcal{H}' \simeq -\frac{1}{2}\mathcal{H}^2$ that hold well inside the matter dominated epoch. The propagation speeds for the high frequency modes with $k^2 \gg \max\{\mathcal{H}^2, R\Gamma\mathcal{H}\}$ are not modified by the interaction so we still have the vanishing sound speed of the dust component and $c_{\rm s}$ for the dark energy fluid in that regime. Thus, no risk of UV Laplacian instabilities arise from the interaction, i.e., the characteristic frequencies of the modes will be real and no exponentially growing modes arise.

On the other hand, other instabilities such as tachyonic modes could be triggered by the interaction due to its effects on the friction and the effective masses. Since the interaction is completely negligible on super-Hubble scales because of the common large scale frame of all the fluids, infrarred instabilities do not arise. This is easy to understand because the standard non-interacting evolution is recovered at those scales. The interaction however leads to the appearance of a second horizon for the dark energy fluid given that is given by the relation

$$k_\star^2 = \left| 3\left(1 - \frac{w}{c_{\rm s}^2}\right) R\Gamma\mathcal{H} \right| . \tag{B.2}$$

This scale is obtained by comparing the two competing terms $c_{\rm s}^2 k^2$ and $3(c_{\rm s}^2 - w)\mathcal{H}R\Gamma$ in the second line of (B.1). For sub-Hubble scales that are above this horizon $\mathcal{H} \ll k \ll k_\star$, the k-dependence of the equations disappear because the only $k-$dependent term $c_{\rm s}^2 k^2$ can be safely neglected and, if we consider the strongly coupled regime[4] $\Gamma \gg \mathcal{H}$, the determinant of

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[4]Notice that there is a wide band of modes in this regime since $k_\star \gg \mathcal{H}$.

the mass matrix $\hat{M}$ in the second line of (B.1) reduces to

$$\det \hat{M} \simeq -\frac{9}{2}(c_s^2 - w)R\Gamma\mathcal{H}^3 \tag{B.3}$$

From this mass term we expect to have a negative eigenvalue corresponding to the unstable mode leading to the Jeans collapse of the DM component so in order to avoid an additional tachyonic mode we need to impose $\det \hat{M} < 0$ that is achieved for $\Gamma > 0$, i.e., $\alpha > 0$. One could argue why would we want to remove the growing modes, especially since we already have the usual Jeans instability. The reason is that the growing modes associated to the interaction will develop on a time scale parameterically determined by $R\Gamma$ that will be much shorter than the cosmological Hubble time scale since we are in the regime $\Gamma \gg \mathcal{H}$. We do not intend to perform a detailed analytical analysis of the evolution of the perturbations here, but simply motivate why we mostly restrict to $\alpha > 0$ in our analysis since this guarantees $\Gamma > 0$ and, therefore, the avoidance of the discussed instabilities. Furthermore, we have checked also numerically that the region $\alpha < 0$ typically leads to instabilities that spoil the possibility of sensible cosmologies unless $|\alpha|$ is sufficiently small.

C Survey specifications for bins and galaxy densities

In the following tables we show the redshift bins distributions and the densities of Luminous Red Galaxies (LRG), Emission Line Galaxies (ELG), Quasars (QSO) and Bright Galaxies (BGS) for the differents surveys used, both for clustering (J-PAS, DESI and spectroscopic Euclid) and lensing analysis (J-PAS and photometric Euclid). Galaxy densities are in units of $10^{-5}h^3\,\mathrm{Mpc}^{-3}$.

J-PAS
z	LRG	ELG	QSO
0.3	226.6	2958.6	0.45
0.5	156.3	1181.1	1.14
0.7	68.8	502.1	1.61
0.9	12.0	138.0	2.27
1.1	0.9	41.2	2.86
1.3	0	6.7	3.60
1.5	0	0	3.60
1.7	0	0	3.21
1.9	0	0	2.86
2.1	0	0	2.55
2.3	0	0	2.27
2.5	0	0	2.03
2.7	0	0	1.81
2.9	0	0	1.61
3.1	0	0	1.43
3.3	0	0	1.28
3.5	0	0	1.14
3.7	0	0	0.91
3.9	0	0	0.72

DESI				
z	BGS	LRG	ELG	QSO
0.1	2240	0	0	0
0.3	240	0	0	0
0.5	6.3	0	0	0
0.7	0	48.7	69.1	2.75
0.9	0	19.1	81.9	2.60
1.1	0	1.18	47.7	2.55
1.3	0	0	28.2	2.50
1.5	0	0	11.2	2.40
1.7	0	0	1.68	2.30

Euclid_ph	
z	ELG
0.3	7440
0.5	6440
0.7	5150
0.9	3830
1.1	2670
1.3	1740
1.5	1070
1.7	620
1.9	341
2.1	178
2.3	88.3
2.5	41.8

Euclid_sp	
z	ELG
1.0	68.6
1.2	55.8
1.4	42.1
1.6	26.1

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