

MULTIWAVELENGTH OBSERVATIONS OF MASTER OT 075353.88+174907.6: A LIKELY SUPEROUTBURST OF A LONG PERIOD DWARF NOVA SYSTEM

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ABSTRACT

MASTER OT 075353.88+174907.6 was a blue optical transient reported by the MASTER-Net project on 2017 Oct 31. We carried out multiwavelength follow-up observations of this source during its 2017 outburst using *Swift* and RATIR. The source was found to be $\gtrsim 4.4$ mag above its quiescent level during the peak of the outburst and the outburst lasted $\gtrsim 19$ days. Our observations suggest that it was a superoutburst of a long period U Geminorum type dwarf nova system. The spectral energy distribution during the initial slow decay phase of the outburst was consistent with a disk-dominated spectrum (with spectral indices $\Gamma \approx 1.5\text{--}2.3$). Afterwards, the UV flux decreased slower than the optical and the spectral energy distribution was very steep with indices $\Gamma \approx 3.7 \pm 0.7$. This slow UV decay may be the emission from a cooling white dwarf heated during the outburst. No X-ray emission was detected from the source since it is likely located at a large distance > 2.3 kpc.

RESUMEN

MASTER OT 075353.88+174907.6 fue un objeto transitorio reportado por el proyecto MASTER-Net en 2017 Oct 31. Realizamos una campaña multi-onda de este objeto durante su erupción en 2017 con *Swift* y RATIR. La fuente fue descubierta $\gtrsim 4.4$ mag por encima de su nivel mínimo en el máximo de la erupción, la cual duró $\gtrsim 19$ días. Nuestras observaciones sugieren que se trata de una supererupción de un sistema tipo U Geminorum de largo período. La distribución espectral de energía en la fase inicial de la caída es consistente con un espectro dominado por el disco de acreción (con índices espectrales $\Gamma \approx 1.5\text{--}2.3$). Después, el flujo UV decreció más lentamente que el óptico, con una distribución espectral de energía con índices $\Gamma \approx 3.7 \pm 0.7$. El lento declive en el UV puede deberse al enfriamiento de la enana blanca, la cual fue calentada durante la erupción. No se detectó emisión de la fuente en rayos-X, lo cual sugiere una distancia mucho mayor que 2.3 kpc.

Key Words: accretion, accretion discs — novae, cataclysmic variables — stars: dwarf novae — stars: individual: MASTER OT 075353.88+174907.6 — ultraviolet: stars

1. INTRODUCTION

Dwarf novae are a sub-set of cataclysmic variables — binary systems hosting a white dwarf and a low-mass main sequence star. In these systems, the companion is transferring mass to the white dwarf via Roche-lobe overflow and an accretion disk is

formed around it. Dwarf novae can undergo outbursts, thought to occur due to thermal instabilities in the disk (see Lasota 2001, for details) resulting in a brightness increase of $\approx 6\text{--}100$ times than its quiescence level (Osaki 1996). While these dwarf novae outbursts can be very homogeneous from a single source, occasionally they exhibit ‘superoutbursts’. These are different from normal outbursts as they are brighter and can last 5–10 times longer

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(≈ 15 – 40 days compared to ≈ 2 – 20 days in normal outbursts, the exact values depend on the type of system; Mauche et al. 2001).

MASTER OT 075353.88+174907.6 (hereafter OT 0753) was reported as a optical transient on 2017 October 31 (Balanutsa et al. 2017) by the MASTER-Net project (Kornilov et al. 2012). This transient had a B magnitude of 18.6 during the discovery observation. The Sloan Digital Sky Survey (SDSS) had previously observed the source position of OT 0753 and did not detect any source. The sensitivity limit of the SDSS in the g' band (which is the closest one to the Johnson B band from the MASTER-Net project) is ≈ 23 mag. This indicates (see also Balanutsa et al. 2017) that during the MASTER-Net detection OT 0753 was $\gtrsim 4.4$ mag brighter than during its known quiescent level. Furthermore, a NUV detection within 5 arcsec of this source position was found in the *Galaxy Evolution Explorer* mission (*GALEX*; Martin et al. 2005) database. This source was observed on 2006 February 7 using *GALEX* and during that observation it had a magnitude of ≈ 22.9 in the NUV (Balanutsa et al. 2017). The source was not detected in the *GALEX*/FUV band. The *GALEX*/NUV detection and SDSS non-detection suggest that the source hosts a blue object when in quiescence.

To investigate the nature of this source, we carried out a multiwavelength campaign across the near-infrared, optical, UV, and X-ray bands during its 2017 outburst. We discuss the results of our observing campaign and show that this blue optical transient is likely a superoutburst of a dwarf nova which hosts a hot white dwarf that dominates the emission in quiescence.

2. OBSERVATIONS AND DATA ANALYSIS

We requested follow-up observations of OT 0753 using the *Neil Gehrels Swift Observatory* (Gehrels et al. 2004). We observed OT 0753 a total of 9 times using *Swift* (see Table 1 for a log of these observations). We obtained the first observation within a day of the initial report of the transient (see also Parikh & Wijnands 2017). This observation was carried out over all six optical and UV bands using the Ultraviolet and Optical Telescope (UVOT; Roming et al. 2005). Subsequent observations using the UVOT only observed the source in the three UV bands. Further UV and X-ray data of the source were obtained using *Swift* up to ≈ 1 month after the initial detection, until the source had approached its assumed quiescent level.

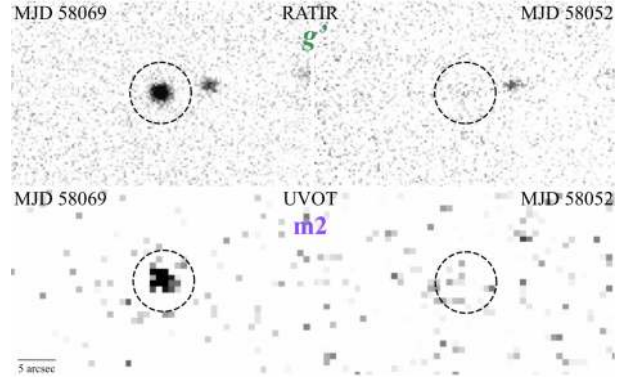


Fig. 1. The detection of OT 0753 (indicated by the dashed black circle) in outburst and its subsequent quiescence phase is shown. Top two panels: the RATIR images in the optical g' band; bottom two panels: the *Swift*/UVOT images in the UV $m2$ band. The color figure can be viewed online.

We analysed all the *Swift*/UVOT data using the Level 2 products. We added all the individual exposures (in the same band) of a given observation using `uvotimsum`. We selected a circular source extraction region having a radius of 5 arcsec. The background used was a circular region having a radius of 10 arcsec, placed on a source-free location on the CCD. The source magnitude and flux were extracted using the `uvotsource` tool. The results of this analysis is shown in Table 1.

The source was not detected in X-rays throughout our observing campaign. The upper limits on the source count rate (in the 0.5–10 keV energy range) during the individual observations taken with the X-ray Telescope (XRT; Burrows et al. 2005) observations were $\lesssim (0.8 - 2.4) \times 10^{-3}$ counts s^{-1} (depending on the exposure time of each observation; calculated using the 90% prescription by Gehrels 1986). To obtain the flux upper limits, we simulated X-ray spectrum using `Xspec` (Arnaud 1996) employing an absorbed power-law model and assuming that a power-law index of 2 was representative of the spectrum (similar to what has been observed from cataclysmic variables that were detected in the X-rays; Done & Osborne 1997; Balman 2015). The equivalent hydrogen column density N_H was obtained using the HEASARC N_H tool³ and was found to be 3.8×10^{20} cm^{-2} . Our obtained count rate upper limits corresponded to flux upper limits of $\lesssim (4 - 9) \times 10^{-14}$ erg cm^{-2} s^{-1} in the 0.5–10 keV

³This was done using the HEASARC N_H tool: <https://heasarc.gsfc.nasa.gov/cgi-bin/Tools/w3nh/w3nh.pl>; (Dickey & Lockman 1990).

range. We stacked all the XRT observations using `ximage` to check whether the source could be detected if we used all available data. We found that the source was still not detected in ≈ 13.4 ksec of stacked XRT images and we obtained a stricter X-ray upper limit of $\lesssim 1.7 \times 10^{-4}$ counts s^{-1} (corresponding to a flux upper limit of $\lesssim 6 \times 10^{-15}$ erg $cm^{-2} s^{-1}$, again assuming a photon index of 2).

The optical and near-infrared follow-up was carried out by a ground-based campaign using the 1.5 m RATIR (Reionization and Transients InfraRed) telescope located at the Mexican Observatorio Astronómico Nacional (OAN) on the Sierra San Pedro Mártir (SPM) in Baja California (Watson et al. 2012). This campaign lasted from 2017 November 12 to 2017 November 25 and included 12 nights of observations with exposure times ranging from ≈ 0.8 –2.4 hours per night (see Table 1 for a log of the RATIR observations). On each night, the source was observed using the SDSS g' , i' , and z' bands. We analysed the RATIR data using the PYTHON package `PhotoPipe`⁴. This pipeline automatically reduced all the data, carrying out the following processing steps: (1) calibrating the images using the bias, dark, and flat fields, (2) performing astrometry, and (3) co-adding the data and performing absolute photometry. For both the RATIR as well as the UVOT results, all errors are given for the 1σ confidence level and all magnitudes correspond to the AB system. The light curve constructed using these data is shown in Figure 2.

3. RESULTS

The first notice of the outburst of OT 0753 was published by the MASTER-Net Project on 2017 October 31 (in the B band, indicated by the blue diamond in the upper panel of Figure 2; Balanutsa et al. 2017). The rise to outburst of OT 0753 was not detected and the exact day on which the outburst began is not known. We examined the *Gaia* alerts⁵ and found that it was detected by *Gaia* on 2017 October 28 (≈ 3 days before the reported detection from the MASTER-Net project; the source was observed twice using the *Gaia* G broadband filter, indicated by the blue-green squares in the upper panel of Figure 2). However, this information was not immediately publicly available and could only be retrospectively obtained. We examined the pre-outburst *GALEX* and SDSS data for this source. We assumed that the archival *GALEX* data (detected in the NUV and not detected in the FUV) and the non-

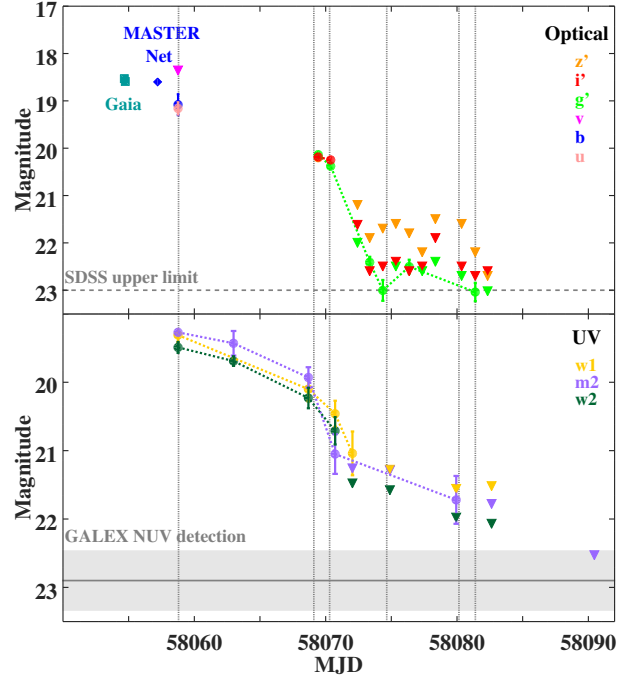


Fig. 2. The evolution of OT 0753 across the various optical/near-infrared (upper panel) and UV bands (lower panel) is shown. The upper panel also shows the detection of OT 0753 as reported by the *Gaia* Photometric Alert (blue-green squares) and the MASTER-Net project (blue diamond; Balanutsa et al. 2017). The SDSS detection limit is shown in the upper panel and the *GALEX*/NUV detection level (along with its error bar as indicated by the grey band) is shown in the lower panel. All magnitudes correspond to the AB system. The grey vertical dashed lines indicate the times for which we have constructed the SEDs shown in Figure 3. The color figure can be viewed online.

detection by the SDSS are representative of our system in (pre-outburst) quiescence (see also Balanutsa et al. 2017). These data are summarized in Table 1. We estimated the upper limit of the source magnitude in the *GALEX*/FUV band using the known sensitivity versus exposure time relation⁶.

The RATIR and *Swift*/UVOT light curves are shown in Figure 2. The outburst displayed a slowly decaying phase that lasted $\gtrsim 16$ days from MJD 58054 to 58070. The first RATIR observation was only obtained near the end of this slow decay phase. OT 0753 was only detected in the i' and z' near-infrared bands using RATIR during this slowly decaying outburst phase. After MJD 58070 the optical magnitude rapidly decayed, by a factor

⁴<https://github.com/maxperry/photometrypipeline>.

⁵<http://gsaweb.ast.cam.ac.uk/alerts/alertsindex>.

⁶http://www.galex.caltech.edu/DATA/gr1_docs/GR1_Observers_guide_v1.htm.

TABLE 1
LOG OF OBSERVATIONS OF OT 0753 WITH *SWIFT*/UVOT AND RATIR.*

<i>Swift</i> /UVOT														
MJD	Observation ID	Total exposure (s)	Magnitude (AB)						Flux ($\times 10^{-17}$ erg s $^{-1}$ cm $^{-2}$ Å $^{-1}$)					
			v	b	u	w1	m2	w2	v	b	u	w1	m2	w2
58058.8	00010373001	988	<18.36	19.1±0.2	19.2±0.1	19.3±0.1	19.3±0.1	19.5±0.1	<16.8	13.4±2.7	19.5±2.4	43.0±2.1	43.0±2.1	42.1±3.0
58063.0	00010373002	1788	–	–	–	–	19.4±0.2	19.7±0.1	–	–	–	–	37.1±6.1	35.0±2.2
58068.7	00010373003	1982	–	–	–	20.1±0.2	19.9±0.2	20.2±0.2	–	–	–	14.6±2.1	23.5±3.3	21.4±2.9
58070.7	00010373004	951	–	–	–	20.5±0.2	21.1±0.3	20.7±0.2	–	–	–	10.4±1.9	8.4±2.3	13.7±2.5
58072.0	00010373005	1033	–	–	–	21.0±0.3	<21.3	<21.5	–	–	–	6.12±1.8	<6.91	<6.8
58075.0	00010373006	1627	–	–	–	<21.3	<21.3	<21.6	–	–	–	<4.93	<6.71	<6.1
58079.9	00010373008	946	–	–	–	<21.6	21.7±0.4	<22.0	–	–	–	<3.81	4.5±1.5	<4.3
58082.6	00010373009	1062	–	–	–	<21.5	<21.8	<22.1	–	–	–	<3.93	<4.3	<3.9
58090.5	00010373010	2948	–	–	–	–	<22.5	–	–	–	–	–	<2.2	–

RATIR								
MJD	Total exposure (s)	Magnitude (AB)			Flux ($\times 10^{-18}$ erg s $^{-1}$ cm $^{-2}$ Å $^{-1}$)			
		g'	i'	z'	g'	i'	z'	
58069.4	8580	20.1±0.1	20.2±0.02	20.2±0.1	42.2±1.8	15.8±0.49	10.8±0.8	
58070.4	3480	20.4±0.1	20.3±0.1	–	33.7±2.2	14.9±1.1	–	
58072.4	3900	<22	<21.62	<21.2	<7.6	<4.2	<4.3	
58073.4	3000	22.4±0.1	<22.6	<21.9	5.2±0.5	<1.7	<2.3	
58074.4	3000	23.0±0.2	<22.5	<21.7	3.0±0.5	<1.9	<2.7	
58075.3	3000	<22.5	<22.4	<21.6	4.8±0.0	<2.1	<3.0	
58076.4	3000	22.5±0.2	<22.6	<21.8	4.8±0.6	<1.7	<2.5	
58077.4	3000	<22.6	<22.5	<22.2	<4.4	<1.9	<1.7	
58078.4	3000	<22.4	<21.9	<21.5	<5.3	<3.3	<3.3	
58080.4	3000	<22.7	<22.5	<21.6	<4.0	<1.9	<3.0	
58081.4	3000	23.0±0.2	<22.7	<22.2	2.9±0.5	<1.6	<1.7	
58082.3	3000	<23.0	<22.6	<22.7	<3.0	<1.7	<1.1	

Summary of observations using other instruments					
Instrument	MJD	Band	Magnitude	Flux ($\times 10^{-18}$ erg s $^{-1}$ cm $^{-2}$ Å $^{-1}$)	
SDSS	–	g'	$\lesssim 23$	$\lesssim 3.0$	
<i>GALEX</i>	53773.4	NUV	22.9±0.5	14.3±8.3	
<i>GALEX</i>	53773.4	FUV	$\lesssim 21.5$	$\lesssim 117.1$	
<i>Gaia</i>	58054.7	G	≈ 18.5	≈ 93.1	
<i>Gaia</i>	58054.8	G	≈ 18.6	≈ 88.1	
MASTER-Net	58057.2	B	≈ 18.6	≈ 14.8	

*The log of the observations of OT 0753 of our *Swift*/UVOT and RATIR observing campaign is tabulated. The outburst observations obtained by *Gaia* and the MASTER-Net project and the archival quiescent observations using SDSS and *GALEX* are also shown. All the errors are presented for the 1σ confidence level and all magnitudes correspond to the AB system.

≈ 3 in ≈ 3 –5 days, to values close to the SDSS upper limit determined when OT 0753 was (assumed to be) in quiescence. Around this time the UV magnitude also began to decrease but this decrease was slow compared to the optical decay. The UV magnitude only decreased by a factor of $\gtrsim 1.5$ over the same time (as compared to a drop by a factor ≈ 3 in the optical in ≈ 3 –5 days). In spite of this decrease, during the next observations (at MJD 58075) the UV magnitude did not approach its assumed pre-outburst quiescent level and continued to decay slowly. The last observation of the source, carried out on MJD 58091, indicated an upper limit of $\lesssim 22.5$ mag in the m2 band. This band is similar to the NUV band on board *GALEX*. Thus, our last

UVOT observation of OT 0753 indicated an upper limit consistent with the known *GALEX*/NUV quiescent level (of $\approx 22.9 \pm 0.5$ mag; see also Figure 2), suggesting that the source was (likely) back in quiescence.

We constructed several spectral energy distributions (SEDs; see Figure 3) of OT 0753 to understand the broadband evolution of the source during its outburst. The (quasi-)simultaneous SEDs were constructed at the times indicated by the vertical dotted grey lines in Figure 2. We calculated the spectral indices (Γ) corresponding to these SEDs using the relation

$$F_\lambda \propto \lambda^{-\Gamma}. \quad (1)$$

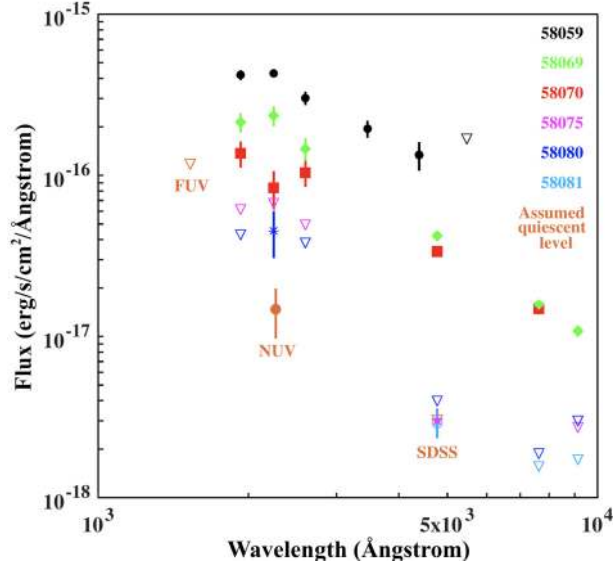


Fig. 3. The SED evolution of OT 0753 at various times is shown. The vertical dotted grey lines in Figure 2 indicate the times for which these SEDs were constructed. The assumed quiescent level using the pre-outburst *GALEX* and SDSS data is shown in orange. The color figure can be viewed online.

The SEDs corresponding to the broadband coverage of the initial slow outburst decay (over MJD 58059 to 58070) are shown in Figure 3 (as black ●, green ◆, and red ■, respectively; upper limits are always shown using ▽ in the appropriate colour). These SEDs fit well a blackbody model. The spectral indices corresponding to these SEDs were $\Gamma \approx 1.5$ – 2.3 and such values are consistent with the expected emission from a disk-like spectrum (Frank et al. 2002).

Several observations during the subsequent rapid outburst decay stage (after MJD 58070) yielded only upper limits with very few detections across the various bands. The SED constructed for the observations around MJD 58075 is unconstrained with $\Gamma \lesssim 4.1$ (plotted as magenta ★ in Figure 3).

Further, we examined the SED shape around MJD 58080 and 58081 (near the end of the rapid optical and corresponding slow UV decay phase, as shown by the dark blue * and light blue ○, respectively, in Figure 3). Unfortunately, detections were only obtained in either the UV *or* the optical bands for each of these days. However, the source magnitude level at this stage of the outburst is not expected to vary much across time scales of day.⁷ Thus,

⁷Our source does not exhibit rebrightenings (see discussion in § 4.3). Furthermore, photometric studies of sources that do not exhibit rebrightenings during this stage of their outburst

we have a quasi-simultaneous detection in the optical and UV bands which suggests a very steep SED, with $\Gamma \approx 3.7 \pm 0.7$.

The discussion in § 4 shows that this very steep SED a few days after the end of its outburst might not be unrealistic for the quiescent state of OT 0753. Fitting a blackbody model to this SED indicated a temperature of $\gtrsim 20,000$ K. The actual temperature may also be indicative of significantly higher temperatures since for an SED corresponding to these high spectral indices our coverage over the optical and UV bands only probes the tail of the Rayleigh-Jeans distribution ($F_\lambda \propto \lambda^{-4}$). A further increase in this temperature does not significantly change the spectral index calculated over the wavelength range we probe.

If the pre-outburst flux is indeed representative of the quiescent flux level of the source, then the source flux must drop further after MJD 58080–58081. This is confirmed by the UV upper limit we obtained around MJD 58091 (see Table 1 and Figure 2) which indeed demonstrates that the source decreased further (at least in the UV).

We also constructed an SED at the assumed pre-outburst quiescent level, using the *GALEX* and SDSS data, as is shown in orange in Figure 3. The spectral index obtained for our assumed pre-outburst quiescent level is $5.3 \gtrsim \Gamma \gtrsim 2.5$. The upper limit determined using the *GALEX*/FUV band is unconstraining. We fit this pre-outburst quiescent level SED with a blackbody model which results in a blackbody temperature $\gtrsim 12,000$ K.

4. DISCUSSION

We report on the near-infrared, optical, and UV behavior of the transient source OT 0753 during its 2017 outburst and subsequent quiescence. The broadband spectral behavior of the source, the length of the outburst ($\gtrsim 19$ days), and the increase in magnitude above its quiescent level ($\gtrsim 4.4$ mag) indicates that this is likely a superoutburst of a dwarf nova.

Otulakowska-Hypka et al. (2016) have presented the statistical properties of several dwarf nova by studying their outbursts. We use their study to learn more about OT 0753 based on its properties that we observe. We note that their study is based only on optical data and therefore we will only use properties determined from our optical observations to compare them to their results. The duration of the superoutburst of OT 0753 in the optical was $\gtrsim 19$ days. From

(when they are returning to quiescence) do not show large variations on time scales of a day (Cannizzo 2012).

this value we can infer (using the trend they observe of the duration of the superoutburst versus the amplitude of the superoutburst; see their Figure 16) an expected optical amplitude increase of $\gtrsim 4$ mag which is consistent with what we have observed for OT 0753. This supports the inference that OT 0753 is a superoutburst of a dwarf nova.

The 2017 superoutburst is the first reported outburst from this system. Using the relationship between the amplitude of the superoutburst and the amplitude of the normal outburst presented by Otulakowska-Hypka et al. (2016, see their Figure 13), we find that the amplitude of the normal outburst is ≈ 1 – 1.5 magnitude fainter than the superoutburst. This indicates that normal outbursts of OT 0753 will only have a magnitude of ≈ 20 in the optical bands meaning that the peak amplitude of such normal outbursts from this source may not be detectable by the various sky surveys. For example, the survey limit of the MASTER-Net project (based on its individual snapshots) is ≈ 20 – 21 mag and that of the Catalina Real Time Survey (Drake et al. 2009), which also observes this part of the sky, is only ≈ 19 – 19.5 mag. This could explain why the superoutburst of this source was the first activity to be detected and reported and why any previous normal outbursts were missed.

The relation between the amplitude of the superoutburst and the recurrence time between two consecutive superoutbursts presented by Otulakowska-Hypka et al. (2016, see their Figure 9) predicts that superoutbursts in this source should recur every ≈ 230 days. It is unknown why the source has not been observed before its 2017 superoutburst. During a superoutburst, similar to the 2017 one we study, the source is above the survey limit of several all sky surveys only for $\lesssim 5$ days. Since the observation cadence of the source location by the various sky surveys is relatively sparse and since the number of days on which the source would be detected is small, it is not very surprising that the source has not been previously detected during a superoutburst. Alternatively, it could also be that OT 0753 has a superoutburst recurrence time longer than the ≈ 230 days predicted using the relation presented by Otulakowska-Hypka et al. (2016). Such a longer recurrence time is observed for sources that have a lower mass transfer rate (e.g., WZ Sge; Lasota et al. 1995).

4.1. Cooling of the White Dwarf?

Observations of white dwarfs after the end of dwarf nova outbursts show that they may be heated

during the outburst and cool once the outburst ceases. White dwarfs in both short and long orbital period systems were found to have cooled by $\approx 4,000$ – $7,000$ K, ≈ 40 – 70 days after the end of their outbursts (Gänsicke & Beuermann 1996; Long et al. 1994; Godon et al. 2017).

We examined OT 0753 during its 2017 outburst and subsequent quiescence. We found that the SED evolves from resembling emission from a hot accretion disk to one that has a very steep spectral index (of $\Gamma \approx 3.7 \pm 0.7$, indicating a blackbody temperature of $\gtrsim 20,000$ K; see also § 3) when the source transitions from the slow decay during the outburst to quiescence. During this transition, the optical magnitude of OT 0753 fell off faster than the UV magnitude, as shown in Figure 2. The abrupt drop in the optical magnitude is likely representative of the cessation of the outburst and the retreat of the disk. The slow decay in the UV magnitude may be the white dwarf cooling after the superoutburst during which it may have been significantly heated due to the accretion. This suggests that the post-outburst spectrum may be dominated by a hot white dwarf. However, it is unknown if the quiescent disk provides a large contribution to this slowly decaying UV flux.

OT 0753 could be a promising source to study cooling in white dwarfs after being heated during its outburst (see discussion above). However, from its position in the sky and its peak outburst flux it seems to be located in the Galactic halo and it may not be close enough (see discussion in § 4.5 for a distance estimate) to allow a sensitive study of this possible cooling.

4.2. OT 0753 Hosts a Hot White Dwarf in Quiescence

Before its 2017 outburst, OT 0753 had only been detected in the NUV band on board the *GALEX* satellite. The source was not detected by the SDSS. The assumed quiescent level inferred using the *GALEX* and SDSS data also exhibits a steep spectral index in quiescence (with $5.3 \gtrsim \Gamma \gtrsim 2.5$). Such a blue quiescent spectrum is likely indicative of the tail of a Rayleigh-Jeans distribution and further supports that OT 0753 may host a hot white dwarf in quiescence even after all the heat deposited on the star during the outburst has been radiated away.

4.3. Orbital Period of OT 0753

The length of a superoutburst is not enough to definitively infer the orbital period of a dwarf nova system, because there exists a degeneracy between

the length of the superoutburst and the orbital period as both systems having a short and long (above the gap) orbital period can exhibit superoutbursts of similar duration. This can be seen from Figure 18 of Otulakowska-Hypka et al. (2016). The effective white dwarf temperature can help break this degeneracy since systems located above the period gap host hot white dwarfs (likely due to the high mass accretion rate on to the white dwarf; see Figure 21 of Pala et al. 2017). Studying the SED we find that OT 0753 likely hosts a hot white dwarf that has an effective temperature $\gtrsim 20,000$ K (see § 3). From Figure 21 in Pala et al. (2017) it can be seen that only one short period system has an effective temperature $>20,000$ K. Thus, this suggests that our source is likely a long period system having a period of ≈ 4 –5 hours.

Several dwarf nova sources exhibit ‘rebrightenings’ which are episodes of increase of flux after the end of the initial outburst (see for e.g., WZ Sge; Patterson et al. 2002; Kato et al. 2009). Their peak fluxes are observed to reach the flux level observed at the end of the slow decay stage of the outburst. These rebrightenings are hypothesized to be caused by a 3:1 resonance, by the interaction between a precessing eccentric accretion disk and the secondary star in certain specific system configurations (O’Donoghue 2000). These systems have a mass ratio that is small enough ($q = M_2/M_1 < 0.25$, where M_1 is the mass of the white dwarf and M_2 is the mass of the donor; Whitehurst 1988) to accommodate a large accretion disk where a 3:1 resonance can occur (O’Donoghue 2000). It is found that these are systems that have short orbital periods ($\lesssim 2.4$ hours). OT 0753 was observed almost every day and showed a continuous decay trend towards the end of its outburst. The subsequent detection and upper limits indicated that likely no rebrightenings occurred and none have been missed by our coverage. This evidence further supports our inference that OT 0753 is likely a source above the period gap and not a short period system. Alternatively, it could also be a short period dwarf nova that experiences a type D outburst. Type D outbursts are outbursts in short period dwarf nova systems that do not exhibit ‘rebrightenings’ because of their specific configuration (see Kato et al. 2009, for details). However, further evidence presented in the discussion (such as evidence of the steep blue spectrum in quiescence dominated by emission from a hot white dwarf) reinforces the suggestion that OT 0753 belongs to the sample of dwarf nova sources that are above the period gap.

4.4. OT 0753 is a U Gem Like System

The source properties of OT 0753 (the inferred white dwarf temperature, the inferred orbital period, and the SED evolution) indicate that it is very similar to U Geminorum. U Gem was studied using the *Hopkins Ultraviolet Telescope* in the UV and was found to have a very blue quiescent spectrum very soon (≈ 10 days) after the end of one of its outbursts (Long et al. 1993). This hot blue spectrum was dominated by emission from the white dwarf and the inferred temperature of this white dwarf was very high at $\approx 38,000$ K (assuming all the UV light came only from the white dwarf; Long et al. 1993). Similar sources (e.g., UZ Serpentis and SS Aurigae; Lake & Sion 2001) also show a blue white dwarf dominated spectra in quiescence and have orbital periods (≈ 4 hours) similar to that inferred for OT 0753. Thus, OT 0753 is probably also a U Gem type system which is dominated by the hot white dwarf in quiescence.

4.5. No X-ray Detection During the Dwarf Nova Outburst

OT 0753 was not detected in X-rays. Güver et al. (2006) studied U Gem during its outburst and hypothesize that its X-ray activity arises from optically thin plasma close to the white dwarf or from the optically thick boundary layer during the outburst. If all dwarf novae are expected to show X-ray emission, albeit faint, due to similar mechanisms when they are in outbursts, then the lack of any X-ray detection from the source may be evidence that the source is located relatively far away. The X-ray flux from U Gem around the peak of its normal 2002 optical outburst, observed using *Chandra*, was found to be $F_X = 3.3 \times 10^{-11}$ erg cm $^{-2}$ s $^{-1}$ (0.8–7.5 keV). U Gem is located at a distance of around 100 pc (Harrison et al. 2004) which corresponds to an outburst luminosity $\approx 4.1 \times 10^{31}$ erg s $^{-1}$. Assuming that our X-ray upper limit determined at the time close to the outburst peak (see § 2) corresponds to the same luminosity as observed for U Gem the distance to OT 0753 would be $\gtrsim 2.3$ kpc. This is a very rough estimate that makes several assumptions (such as all dwarf novae should show X-ray emission at similar luminosities during their outbursts, that the peak X-ray luminosity from a normal outburst of U Gem and the superoutburst of OT 0753 are similar, and that our first XRT observation is representative of the peak of the superoutburst); however, it suggests that OT 0753 is relatively far away.

4.6. *The Insight Obtained from the UV Coverage*

OT 0753 was observed in the optical and UV bands during its outburst and subsequent decay. The UV probes hotter components of the system compared to the optical. For example, it can probe the inner hot accretion flow and, in this case, also the hot white dwarf. We wish to emphasize the importance of UV coverage of dwarf nova outbursts as this coverage can provide insights into the physics of the system which are unavailable when studying the source only in the optical. For example, if we only studied the optical light curve in Figure 2 we would have only been able to conclude that OT 0753 likely experienced a superoutburst which was not followed by any subsequent rebrightenings. The additional UV coverage after the end of the outburst showed that the source magnitude decayed relatively slowly in this band as compared to the optical. This allowed us to infer that the source likely hosts a cooling white dwarf which may have been heated during the preceding accretion outburst. Similarly, the spectrum covering the UV and optical wavelengths gave us more insight than would be inferred by only studying the optical spectral evolution. The optical spectral evolution alone indicated that a strong disk contribution present during the outburst gradually decreased as the outburst transitioned to quiescence. The additional UV coverage indicated that although the disk component faded away the spectrum remained very steep, suggesting the presence of a hot white dwarf. Furthermore, the assumed quiescent spectrum determined using the SDSS optical data and UV data from *GALEX* also exhibited a very steep spectral index which suggested that even after cooling the source hosts a hot white dwarf.

4.6.1. *Further Studies of Cooling White Dwarfs*

Without the accompanying UV coverage we would have missed the potential cooling of the hot white dwarf that OT 0753 hosts. This source is estimated to be at a large distance away (see § 4.5) and is not an excellent candidate to study the cooling in white dwarfs after the end of its accretion outburst. However, similar UV and optical coverage of nearby promising sources that may exhibit cooling will allow us to learn not only about the accretion physics but also to monitor the cooling white dwarf more accurately, thereby allowing us to infer the physics of the white dwarf. Photometry using the *Swift*/UVOT and similar instruments will allow us to sample a large population of cooling white dwarfs and infer thermal properties of the white dwarf.

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REFERENCES

- Arnaud, K. 1996, *ASPC*, 101, 17
 Balanutsa, P., Gress, O., Shumkov, V., et al. 2017, *ATel*, 10915, 1
 Balman, S. 2015, *AcPPP*, 2, 116
 Burrows, D. N., Hill, J., Nousek, J., et al. 2005, *SSRv*, 120, 165
 Cannizzo, J. K. 2012, *ApJ*, 757, 174
 Dickey, J. M. & Lockman, F. J. 1990, *ARA&A*, 28, 215
 Done, C. & Osborne, J. P. 1997, *MNRAS*, 288, 649
 Drake, A. J., Djorgovski, S. G., Mahabal, A., et al. 2009, *ApJ*, 696, 870
 Frank, J., King, A., & Raine, D. J. 2002, *Accretion Power in Astrophysics*, (3rd ed.; Cambridge, UK:CUP)
 Gänsicke, B. T. & Beuermann, K. 1996, *A&A*, 309, L47
 Gehrels, N. 1986, *ApJ*, 303, 336
 Gehrels, N., Chincarini, G., Giommi, P., et al. 2004, *ApJ*, 611, 1005
 Godon, P., Shara, M. M., Sion, E. M., & Zurek, D. 2017, *ApJ*, 850, 146
 Güver, T., Uluyazi, C., Özkan, M. T., & Göğüş, E. 2006, *MNRAS*, 372, 450
 Harrison, T. E., Johnson, J. J., McArthur, B. E., et al. 2004, *AJ*, 127, 460
 Kato, T., Imada, A., Uemura, M., et al. 2009, *PASJ*, 61, S395
 Kornilov, V. G., Lipunov, V. M., Gorbvskoy, E. S., et al. 2012, *ExA*, 33, 173
 Lake, J. & Sion, E. M. 2001, *AJ*, 122, 1632

- Lasota, J.-P. 2001, *NewAR*, 45, 449
- Lasota, J. P., Hameury, J. M., & Hure, J. M. 1995, *A&A*, 302, L29
- Long, K. S., Blair, W. P., Bowers, C. W., et al. 1993, *ApJ*, 405, 327
- Long, K. S., Sion, E. M., Huang, M., & Szkody, P. 1994, *ApJ*, 424, L49
- Martin, D. C., Fanson, J., Schiminovich, D., et al. 2005, *ApJ*, 619, L1
- Mauche, C. P., Mattei, J. A., & Bateson, F. M. 2001, *ASPC* 229, *Evolution of Binary and Multiple Star Systems*, ed. P. Podsiadlowski, S. Rappaport, A. R. King, F. D'Antona, & L. Burder(ASP), 367
- O'Donoghue, D. 2000, *NewAR*, 44, 45
- Osaki, Y. 1996, *PASP*, 108, 39
- Otulakowska-Hypka, M., Olech, A., & Patterson, J. 2016, *MNRAS*, 460, 2526
- Pala, A. F., Gänsicke, B. T., Townsley, D., et al. 2017, *MNRAS*, 466, 2855
- Parikh, A. S. & Wijnands, R. 2017, *ATel*, 10922, 1
- Patterson, J., Masi, G., Richmond, M. W., et al. 2002, *PASP*, 114, 721
- Roming, P. W. A., Kennedy, T. E., Mason, K. O., et al. 2005, *SSRv*, 120, 95
- Watson, A. M., Richer, M. G., Bloom, J. S., et al. 2012, *SPIE*, 8444, *Ground-based and Airborne Telescopes IV*
- Whitehurst, R. 1988, *MNRAS*, 232, 35