

# Non-linear propagation of kink waves to the solar chromosphere

M. Stangalini<sup>1</sup>, F. Giannattasio<sup>1</sup>, and S. Jafarzadeh<sup>2</sup>

<sup>1</sup> INAF–OAR National Institute for Astrophysics, 00040 Monte Porzio Catone (RM), Italy  
e-mail: marco.stangalini@oa-roma.inaf.it

<sup>2</sup> Institute of Theoretical Astrophysics, University of Oslo, PO Box 1029 Blindern, 0315 Oslo, Norway

Received 4 November 2014 / Accepted 25 February 2015

## ABSTRACT

Small-scale magnetic field concentrations (magnetic elements) in the quiet Sun are believed to contribute to the energy budget of the upper layers of the Sun's atmosphere, as they are observed to support a large number of magneto-hydrodynamic modes. In recent years, kink waves in magnetic elements were observed at different heights in the solar atmosphere, from the photosphere to the corona. However, the propagation of these waves has not been fully evaluated. Our aim is to investigate the propagation of kink waves in small magnetic elements in the solar atmosphere. We analysed high-quality, long duration spectropolarimetric data of a photospheric quiet Sun region observed near the disk centre with the spectropolarimeter CRISP at the Swedish Solar Telescope (SST). We complemented these data with simultaneous and co-spatial broadband chromospheric observations of the same region. Our findings reveal a clear upward propagation of kink waves with frequency above 2.6 mHz. Moreover, the signature of a non-linear propagation process is also observed. By comparing photospheric to chromospheric power spectra, no signature of an energy dissipation is found at least at the atmospheric heights at which the data analysed originate. This implies that most of the energy carried by the kink waves (within the frequency range under study <17 mHz) flows to upper layers in the Sun's atmosphere.

**Key words.** Sun: photosphere – Sun: magnetic fields – Sun: oscillations

## 1. Introduction

Magneto-hydrodynamic (MHD) waves are ubiquitous in the solar atmosphere. Both large-scale magnetic structures, extending over several Mm in length, and small-scale magnetic concentrations (magnetic elements) at scales close to the resolution limit of modern solar telescopes (~120–150 km on the solar photosphere), host MHD waves of different forms (e.g. kink, sausage, Alfvén; Roberts & Webb 1978; Spruit 1981; Edwin & Roberts 1983; Roberts 1983; Musielak et al. 1989; Steiner et al. 1998; Hasan et al. 2003; Musielak & Ulmschneider 2003a; Khomenko et al. 2008; Fedun et al. 2011; Vigeesh et al. 2012; Nutto et al. 2012). Since their discovery, these waves were immediately recognized as fundamental to the energy budget of the solar upper atmosphere. This hypothesis became even more attractive after the availability of very high-resolution spectropolarimetric data, which have provided strong observational proofs that magnetic elements also cover a significant fraction of the solar photosphere (Lagg et al. 2010; Bonet et al. 2012).

The ubiquitous although non-homogeneously distributed (Stangalini 2014) magnetic elements are perturbed by the external photospheric plasma flows and are advected in the solar photosphere (Abramenko et al. 2011; Lepreti et al. 2012; Giannattasio et al. 2013, 2014a,b). Recent observations of the photosphere have revealed that different types of waves can co-exist and interact in the same magnetic element (Stangalini et al. 2013b). Jess et al. (2012) reported of propagating compressible waves in magnetic elements with periods of 110–600 s in both the photosphere and the chromosphere. Moreover, Morton et al. (2013) observed torsional modes (Erdélyi & Fedun 2007) generated by vortices in the solar photosphere. These are just two examples demonstrating the large variety of waves that have been observed so far in magnetic elements.

Morton et al. (2014) investigated the propagation of incompressible waves to the chromosphere, finding a good agreement between photospheric and chromospheric velocity power spectra at frequencies below 8 mHz. Among the many MHD waves that magnetic elements can support, kink waves induced by the solar convection have long been considered as a viable mechanism for transferring energy from the photosphere to the upper layers of the Sun's atmosphere (Kalkofen 1997; Hasan et al. 2003; Musielak & Ulmschneider 2003b,a; Hasan & van Ballegoijen 2008). Kink waves have been reported at different layers in the solar atmosphere; from the lower photosphere (e.g. Keys et al. 2011), to the chromosphere (e.g. Jafarzadeh et al. 2013), and the corona (e.g. McIntosh et al. 2011).

Stangalini et al. (2014) observed the presence of subharmonics, with a fundamental period of 7.6 min, consistent with the granular timescale. This constitutes observational evidence for the excitation of kink waves by the solar convection. The same authors also argued that the detected subharmonics might be the signature of chaotic excitation. Indeed, subharmonics represent a general characteristic of complex chaotic systems in the presence of strong forcing (Sander & Yorke 2009, 2010).

Using both simulations and observations, Keys et al. (2011) found that the majority of magnetic elements have horizontal velocities between 0 and 1 km s<sup>-1</sup>, and 6% of them have velocities in excess of 2 km s<sup>-1</sup>. The same authors argued that this significant fraction of magnetic elements may contribute to heating the upper atmospheric layers, since velocities over 2 km s<sup>-1</sup> can efficiently drive kink waves (Choudhuri et al. 1993). Additionally, Jafarzadeh et al. (2013) observed very short duration velocity pulses up to 15 km s<sup>-1</sup> in the inter-network chromospheric magnetic elements that are well in excess of their mean value (~2 km s<sup>-1</sup>), and that may significantly contribute to the energy budget of the upper layers of the Sun.

In addition to the aforementioned detection of subharmonics, [Stangalini et al. \(2013a\)](#) observed the presence of high-frequency peaks of up to 10–12 mHz in the power spectra of kink oscillations in the solar photosphere. These high-frequency oscillations are well above the cut-off frequency expected for kink waves, which is always smaller than the acoustic cutoff at  $\sim 5.3$  mHz ([Spruit 1981](#)). Therefore, they are expected to propagate upwards. Although the physical properties of the magnetic elements analysed were quite similar, [Stangalini et al. \(2013a\)](#) found that each individual magnetic element had its own signature in the power spectrum. This demonstrates that the kink oscillations are largely influenced by the ambient surrounding conditions.

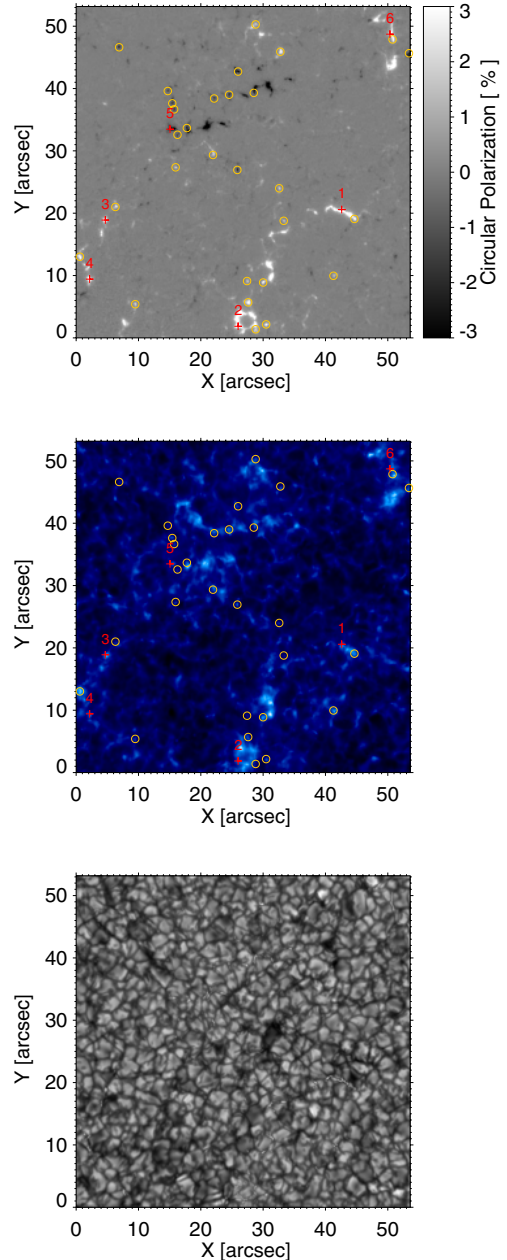
[Morton et al. \(2014\)](#) also studied the average power spectra of kink waves in both the chromosphere and the corona. Their main finding is that a considerable part of energy is dissipated between the chromosphere and the corona itself. This was demonstrated by the observed damping of the coronal power spectra, although the chromospheric and coronal data were not sampled simultaneously and were not co-spatial. However, this study, as well as those reporting the presence of kink oscillations at different heights in the Sun’s atmosphere, suffers a lack of information on the propagation of these waves between different atmospheric layers because of a lack of co-spatial and simultaneous data from different heights in the solar atmosphere.

To overcome this limitation, we analysed high-cadence and high-resolution simultaneous data of the solar photosphere and chromosphere, and investigated the propagation of kink waves in small magnetic elements in the quiet Sun, thus complementing previous results reported in the literature.

## 2. Data set and methods

The data set analysed in this work was acquired on 2011 August 6 with the imager spectropolarimeter CRISP ([Scharmer 2006](#); [Scharmer et al. 2008](#)) at the Swedish Solar Telescope (SST, [Scharmer et al. 2003](#)), and was aimed to study of a quiet Sun region observed at disk centre. It consists of a full-Stokes spectropolarimetric time series of spectral scans acquired at 30 wavelengths in the photospheric Fe I doublet at 630 nm, with a spectral sampling of  $\sim 0.0044$  nm. In addition, simultaneous and co-spatial chromospheric broadband images taken at the core of the Ca II H line at 396.9 nm were acquired. For the spectral lines, the following core-formation heights are assumed: i) 230 km for the Fe I line at 630.1 nm; ii) 200 km for the Fe I line at 630.2 nm ([Faurobert et al. 2012](#)); and iii) 450 km for the Ca II H line at 396.9 nm ([Jafarzadeh et al. 2013](#)). The observations started at 07:57:39 UT and lasted for 47 min with a cadence of the spectral scans of 28 s (100 spectral scans). The pixel scale was  $\sim 0''.059/\text{pixel}$  for the Fe I data and  $0''.034/\text{pixel}$  for the Ca II H data. The Ca II H and Fe I data were co-aligned and remapped to the same FoV of area  $\sim 53'' \times 53''$ . The theoretical full-width half maximum (FWHM) of the point spread function (PSF) at 630 nm is  $0''.16$  or 120 km on the solar photosphere. This constituted the diffraction limit of the telescope. The spectral resolution of CRISP at these wavelengths is 0.0055 nm ([de la Cruz Rodríguez et al. 2015](#)).

The standard CRISP calibration procedure was applied to the data. This includes the CRISPRED calibration pipeline ([de la Cruz Rodríguez et al. 2015](#)), and the MOMFBD (Multi Object Multi Frame Blind Decomposition, [van Noort et al. 2005](#)) restoration aimed at limiting seeing-induced distortions in the images. The photospheric Fe I data were used to estimate the total circular polarization CP within each pixel, which



**Fig. 1.** Example of the data analysed in this study. *Top*: photospheric circular polarization map saturated between  $-3\%$  and  $3\%$ . *Centre*: chromospheric broadband image of the same region. The 35 longest-lived magnetic elements used in the analysis are highlighted with yellow circles. The red crosses mark the six elements whose power spectra are shown in Fig. 3 (see text for more details). *Bottom*: continuum image taken at 395.3 nm with  $1 \text{ \AA}$  bandpass.

is defined as

$$\text{CP} = \int |V(\lambda)| \frac{d\lambda}{I_c}, \quad (1)$$

where  $V(\lambda)$  is the  $V$ -Stokes profile, and  $I_c$  the average continuum intensity.

We remark that CP signals can be attributed to the line-of-sight (LOS) magnetic field. As a consequence, in what follows small-scale CP signals can be identified with magnetic elements. In Fig. 1 we show an example of signed CP polarization map, where the sign is assumed to be that of the blue lobe in Stokes- $V$

profiles, and the corresponding simultaneous broadband image in the core of the Ca II H line. In the same image, we also show a photospheric continuum image of the same field taken at 395.3 nm (1 nm width). Our analysis is based upon the identification and tracking of the longest-lived magnetic elements observed at both the photospheric and chromospheric heights.

The studied magnetic elements were identified and tracked using the well-tested YAFTA code (Welsch & Longcope 2003; DeForest et al. 2007), providing it with the CP maps in input. The parameters of the code were set to track only those magnetic elements i) whose area was larger than  $4 \times 4$  pixel<sup>2</sup> (or equivalently  $0''.24 \times 0''.24$ ), which takes into account seeing-induced adaptive optics (AO) residual aberrations that make the FWHM larger than the theoretical (seeing-free) value; and ii) whose signal in the CP maps was above two standard deviations of the CP image. A total of 185 magnetic elements were found to match these criteria. Among these, we selected the longest lived (at least 90 steps or equivalently  $\sim 42$  min), and whose chromospheric counterpart were clearly identifiable by visual inspection of the simultaneous data. This yielded 35 magnetic elements, which are marked with yellow circles in Fig. 1.

The mean effective radius of the studied magnetic elements, defined as  $\sqrt{\text{Area}/\pi}$ , is  $\langle R_e \rangle \approx 280$  km.

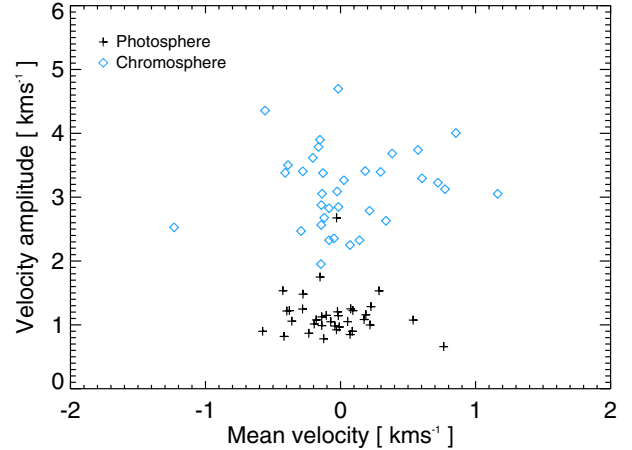
The magnetic elements selected are, on average, co-aligned within  $\sim 2$  pixels. Given the height difference of the sampled atmospheric layers ( $\sim 220$  km), this yields an inclination angle of  $\sim 22$  degrees. This value is in good agreement with Jafarzadeh et al. (2014). We performed tracking of the chromospheric features by applying a procedure based on the search for the centre of mass of the intensity distribution in windows of area  $10 \times 10$  pixel<sup>2</sup>, encompassing the intensity enhancement co-spatial to the photospheric feature. At each time step, the horizontal velocity of each identified magnetic element was computed as the time derivative of the measured position of the feature.

We applied a standard FFT-based analysis to compute the horizontal velocity power spectrum of each photospheric magnetic element and its chromospheric counterpart. In doing this, we only considered one component of the horizontal velocity. This was to avoid frequency doublings due to the square root operator used in the composition of the two components of the horizontal velocity.

### 3. Results

In Fig. 2 we plot the photospheric (black plus symbols) and chromospheric (cyan diamonds) rms velocity amplitude of the magnetic elements as a function of their mean velocity i.e. the mean velocity during the whole lifetime of the element. As we can see, the velocity amplitudes are appreciably larger in the chromosphere than in the photosphere. This can be also seen in Fig. 3, where we show the power spectra of the horizontal velocity for six representative cases among the 35 studied. The chromospheric power spectra are largely amplified with respect to their photospheric counterpart. On top of that, they are also characterized by the presence of different peaks from 2–3 mHz, up to 13–15 mHz. It is worth noting that the amplification is not constant throughout the whole spectral range. Indeed, different spectral bands are more amplified than others. The different magnetic elements do not share the same peaks, and each of them has its own spectral signature.

In Fig. 3, we also show the confidence levels of each pair of spectra obtained from a randomization test.



**Fig. 2.** Chromospheric and photospheric velocity amplitude of the 35 analysed magnetic elements, as a function of their mean velocity.

This test consists in the random shuffling of the measurements of the time series associated with the horizontal velocity of each magnetic element. More in detail, this shuffling is repeated a thousand times for each time series and, at each step, the power spectral density is computed. The number of times a certain peak in the spectrum is larger than that obtained from the original time series (i.e. the correctly ordered one) is counted. This gives the probability that a certain peak in the power spectrum is obtained from a noisy time series, and constitutes the confidence level.

As clearly shown by the results obtained, the confidence level associated with most of the prominent peaks in the power spectra have a very high confidence level (above 0.9). For this reason the peaks shown in the power spectra of Fig. 3 can be regarded as highly trusted.

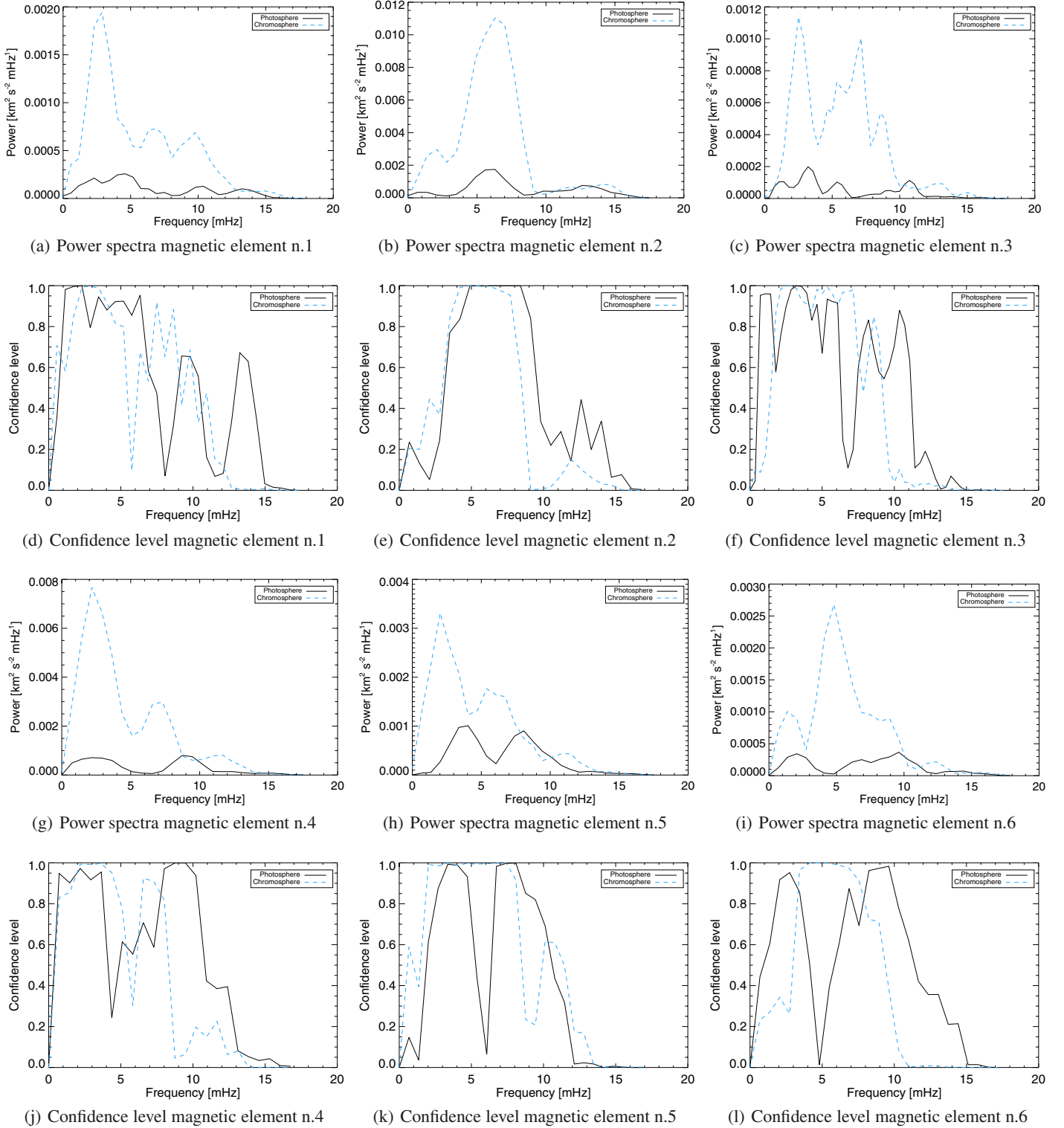
We also studied the wave propagation by applying a phase lag analysis. We estimated the FFT phase and coherence spectra between the photospheric and chromospheric heights for each magnetic element. Only those phase measurements in the spectrum whose coherence level is larger than 0.9 are considered for the estimation of the global phase spectrum obtained from the whole sample. This is necessary to avoid the inclusion of unreliable phase measurements in the final global phase spectrum. We remark that a 0.9 coherence level represents a very restrictive threshold, thus ensuring the reliability of our global phase spectrum estimate.

The results of this analysis are shown in Fig. 4, where we plot the unwrapped global phase spectrum obtained by collecting the most reliable phase measurements (those with coherence larger than 0.9) in the phase spectra among the 35 selected magnetic elements.

In the same figure, we also plot for comparison the phase spectrum expected for propagating kink waves in thin flux tubes as obtained from the dispersion relation of Eq. (4) in Ploner & Solanki (1997). The phase relation is calculated for a propagation speed of  $6 \text{ km s}^{-1}$ , a pressure height scale of  $\sim 130$  km, and a kink cutoff frequency of 2.6 mHz. These are in fact typical values expected for small scale flux tubes (see for example Spruit 1981; Ploner & Solanki 1997). In our sign convention a positive phase difference represents upwards propagation.

We then studied the amplification of the power at chromospheric heights.

The longest-lived 35 magnetic elements were considered to estimate the average power spectrum in both the photosphere

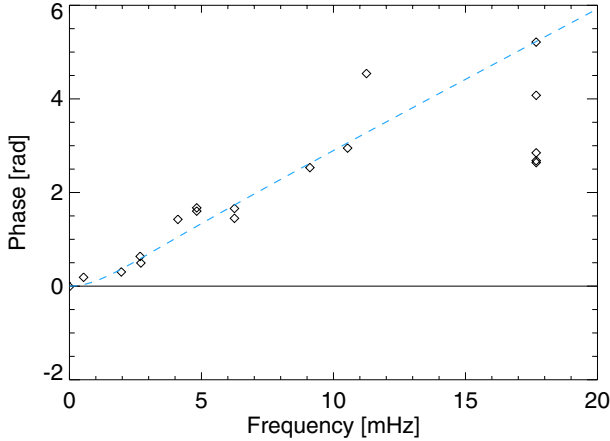


**Fig. 3.** Photospheric and chromospheric power spectra of horizontal velocity oscillations, and confidence level plots obtained from a randomization test of the velocity time series.

and the chromosphere. In Fig. 5 we show the average chromospheric and photospheric spectra (upper panel) and the amplification spectrum (lower panel). The amplitude of kink waves is largely amplified in the 2–8 mHz band, with a maximum around 4.2 mHz where the amplification at the chromospheric heights sampled by the core of Ca II H spectral line reaches a factor of five. The error bars represent the standard error of the mean.

#### 4. Discussions and conclusions

We have studied the propagation of kink waves in magnetic elements between the photosphere and the chromosphere. The power spectra of the horizontal velocity of different magnetic elements in the photosphere show different patterns. This is not surprising since the kink mode is largely influenced by the



**Fig. 4.** Phase diagram obtained from the tracking of the selected magnetic elements (diamonds). Among the phase spectra used and belonging to the selected elements, only the measurements with a coherence above 0.9 are retained. The dashed line represents the expected phase for propagating kink waves at  $6 \text{ km s}^{-1}$  and with a cutoff frequency of 2.6 mHz, between two layers with a separation of 220 km.

ambient plasma surrounding the magnetic elements, which affects their oscillatory behaviour. This result confirms previous findings by [Stangalini et al. \(2013a\)](#), obtained from the analysis of Hinode data at lower spatial and temporal resolution.

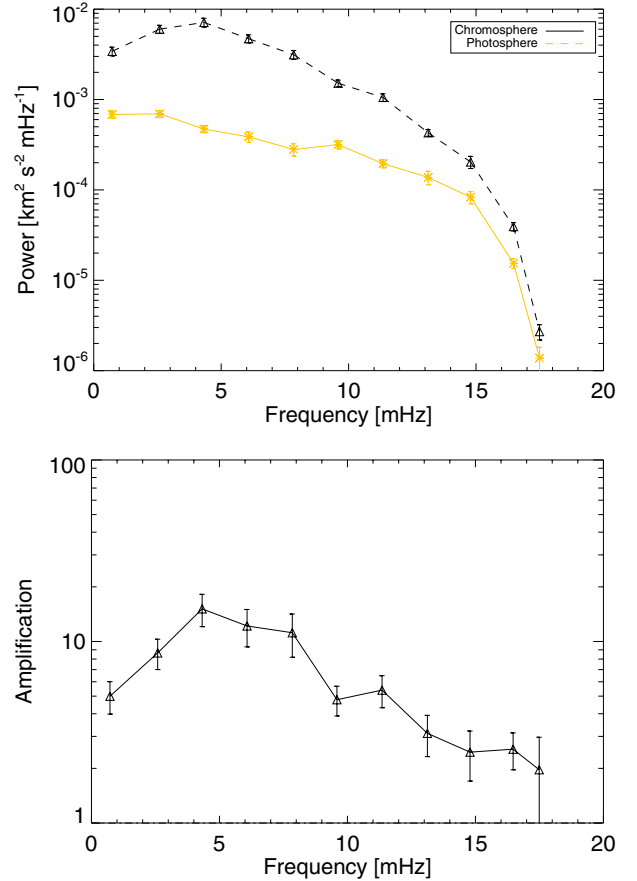
However, the results presented here also reveal, for the first time, that the chromospheric spectra show remarkable differences with respect to their photospheric counterparts. The same magnetic element exhibits a completely different pattern of peaks in the photosphere with respect to the chromospheric power spectrum. This is an evidence for a non-linear wave propagation regime.

It is worthwhile noting here that a randomization test demonstrated the reliability of the obtained power spectra with a confidence level larger than 90%. We remark that the randomization test also validates the quality of the tracking. In more detail, it is unreasonable to find high confidence levels in the spectra of noisy signals. For this reason, the high confidence levels found in the spectra, together with the large coherence found between the two sampled layers, can be regarded as a good validation of the methods employed for the tracking of the magnetic elements at both heights.

We also studied the propagation of the kink waves between the photospheric layer sampled by the Fe I 630.1 nm spectral line and the chromospheric height sampled by the Ca II H spectral data. For this purpose, only the 35 longest-lived magnetic elements were considered to maximize the frequency resolution in the power spectra.

Moreover, considering only frequencies in the phase spectra with very high confidence level between the photosphere and the chromosphere, we estimated the global phase spectrum. This is done to ensure the reliability of the phase estimates themselves. However, we note here that, since the propagation of the kink waves is non-linear, in doing this we are only selecting those cases for which a linear regime can be applied.

The phase diagram so obtained shows the upwards propagation of the kink waves compatible with a velocity of the order of  $6 \text{ km s}^{-1}$ , and a cutoff frequency of 2.6 mHz. Such value for the cutoff frequency is consistent with the theoretical expectations ([Spruit 1981](#)), and represents a first experimental test of this in the lower solar atmosphere.



**Fig. 5.** Top: kink photospheric and chromospheric power spectra obtained from 35 magnetic elements and averaged in bins of 0.8 mHz. Bottom: amplification spectrum.

It is worth remarking here that the inclination of the magnetic elements can significantly affect the theoretical phase relation shown for comparison in Fig. 4. This is because the dispersion relation of kink waves (see [Ploner & Solanki 1997](#)) depends on the acoustic cutoff frequency, which is in turn lowered in inclined magnetic fields (see for example [McIntosh & Jefferies 2006](#)). In our case, the selected magnetic elements have an inclination angle of  $\sim 22$  degrees, thus they can be considered nearly vertical. However, the upwards propagation of kink waves is unambiguously revealed by the phase analysis of Fig. 4. This result is valid independent of the inclination angle and other physical conditions of the flux tubes that can affect, to some extent, the exact values of the cutoff frequency and the propagation speed ([Ploner & Solanki 1997](#)). Our findings show that the chromospheric power spectra are largely amplified. The amplification spectrum shows that most of the enhancement is concentrated in the 1-8 mHz band, although there is no significant power reduction at least up to the limit imposed by the data sampling ( $\sim 17$  mHz).

In the presence of a non-linear propagation regime, kink waves are expected to deposit a significant fraction of energy through dissipation ([Morton et al. 2014](#)). The amplification spectrum estimated from the 35 studied magnetic elements demonstrates that this is not the case at the heights sampled by the data we analysed (up to  $\sim 450$  km above the photosphere). A dissipation of kink waves would in fact result in a deficit of power at chromospheric heights, with respect to photospheric ones. For this reason, we argue that, in absence of deficits of power in the chromospheric spectra, no dissipation should be expected up to

approximately 450 km above the photosphere, where the core of the Ca II H spectral line forms.

Morton et al. (2014) observed the dissipation of kink waves between chromosphere and the corona. In this context, the results presented here demonstrate that the energy contained in the MHD kink waves in small magnetic elements is not dissipated in the low atmosphere but flows entirely to heights that are above the altitude of formation of the core of the Ca II H spectral line. This may have implications for the energy balance of both chromospheric and, in particular, coronal heights, to where most of the energy flows. Finally, the lack of damping is also evidence for the fact that the kink waves do not interact with other modes during the propagation between the heights sampled by our observations.

*Acknowledgements.* This work has been partly supported by the “Progetti di ricerca INAF di Rilevante Interesse Nazionale” (PRIN-INAF) 2010 and PRIN-MIUR 2012 (prot. 2012P2HRCR) entitled “Il sole attivo ed i suoi effetti sul clima dello spazio e della terra” grants, funded by the Italian National Institute for Astrophysics (INAF) and Ministry of Education, Universities and Research (MIUR), respectively. This study received funding from the European Unions Seventh Programme for Research, Technological Development and Demonstration, under the Grant Agreements of the eHEROES (no. 284461, [www.eheroes.eu](http://www.eheroes.eu)), SOLARNET (no. 312495, [www.solarnet-east.eu](http://www.solarnet-east.eu)). We also thank Luc Rouppe van der Voort, and Alice Cristaldi for their help in supporting the data reduction, and Serena Criscuoli for the support during the data acquisition. M.S. and F.G. are grateful to Robertus Erdélyi for the useful discussions and suggestions provided, and to Ilaria Ermolli for a critical reading of the manuscript.

## References

- Abramenko, V. I., Carbone, V., Yurchyshyn, V., et al. 2011, *ApJ*, **743**, 133  
 Bonet, J. A., Cabello, I., & Sánchez Almeida, J. 2012, *A&A*, **539**, A6  
 Choudhuri, A. R., Dikpati, M., & Banerjee, D. 1993, *ApJ*, **413**, 811  
 DeForest, C. E., Hagenaar, H. J., Lamb, D. A., Parnell, C. E., & Welsch, B. T. 2007, *ApJ*, **666**, 576  
 de la Cruz Rodríguez, J., Löfdahl, M. G., Sütterlin, P., Hillberg, T., & Rouppe van der Voort, L. 2015, *A&A*, **573**, A40  
 Edwin, P., & Roberts, B. 1983, *Sol. Phys.*, **88**, 179  
 Erdélyi, R., & Fedun, V. 2007, *Science*, **318**, 1572  
 Faurobert, M., Ricort, G., & Aime, C. 2012, *A&A*, **548**, A80  
 Fedun, V., Shelyag, S., & Erdélyi, R. 2011, *ApJ*, **727**, 17  
 Giannattasio, F., Del Moro, D., Berrilli, F., et al. 2013, *ApJ*, **770**, L36  
 Giannattasio, F., Berrilli, F., Biferale, L., et al. 2014a, *A&A*, **569**, A121  
 Giannattasio, F., Stangalini, M., Berrilli, F., Del Moro, D., & Bellot Rubio, L. 2014b, *ApJ*, **788**, 137  
 Hasan, S. S., & van Ballegoijen, A. A. 2008, *ApJ*, **680**, 1542  
 Hasan, S. S., Kalkofen, W., van Ballegoijen, A. A., & Ulmschneider, P. 2003, *ApJ*, **585**, 1138  
 Jafarzadeh, S., Solanki, S. K., Feller, A., et al. 2013, *A&A*, **549**, A116  
 Jafarzadeh, S., Solanki, S. K., Lagg, A., et al. 2014, *A&A*, **569**, A105  
 Jess, D. B., Shelyag, S., Mathioudakis, M., et al. 2012, *ApJ*, **746**, 183  
 Kalkofen, W. 1997, *ApJ*, **486**, L145  
 Keys, P. H., Mathioudakis, M., Jess, D. B., et al. 2011, *ApJ*, **740**, L40  
 Khomenko, E., Collados, M., & Felipe, T. 2008, *Sol. Phys.*, **251**, 589  
 Lagg, A., Solanki, S. K., Riethmüller, T. L., et al. 2010, *ApJ*, **723**, L164  
 Lepreti, F., Carbone, V., Abramenko, V. I., et al. 2012, *ApJ*, **759**, L17  
 McIntosh, S. W., & Jefferies, S. M. 2006, *ApJ*, **647**, L77  
 McIntosh, S. W., de Pontieu, B., Carlsson, M., et al. 2011, *Nature*, **475**, 477  
 Morton, R. J., Verth, G., Fedun, V., Shelyag, S., & Erdélyi, R. 2013, *ApJ*, **768**, 17  
 Morton, R. J., Verth, G., Hillier, A., & Erdélyi, R. 2014, *ApJ*, **784**, 29  
 Musielak, Z. E., & Ulmschneider, P. 2003a, *A&A*, **400**, 1057  
 Musielak, Z. E., & Ulmschneider, P. 2003b, *A&A*, **406**, 725  
 Musielak, Z. E., Rosner, R., & Ulmschneider, P. 1989, *ApJ*, **337**, 470  
 Nutto, C., Steiner, O., Schaffnerberger, W., & Roth, M. 2012, *A&A*, **538**, A79  
 Ploner, S. R. O., & Solanki, S. K. 1997, *A&A*, **325**, 1199  
 Roberts, B. 1983, *Sol. Phys.*, **87**, 77  
 Roberts, B., & Webb, A. R. 1978, *Sol. Phys.*, **56**, 5  
 Sander, E., & Yorke, J. A. 2009 [[arXiv:0910.3570](https://arxiv.org/abs/0910.3570)]  
 Sander, E., & Yorke, J. A. 2010 [[arXiv:1002.3363](https://arxiv.org/abs/1002.3363)]  
 Scharmer, G. B. 2006, *A&A*, **447**, 1111  
 Scharmer, G. B., Dettori, P. M., Löfdahl, M. G., & Shand, M. 2003, in Innovative Telescopes and Instrumentation for Solar Astrophysics, eds. S. L. Keil, & S. V. Avakyan, SPIE Conf. Ser., 4853, 370  
 Scharmer, G. B., Narayan, G., Hillberg, T., et al. 2008, *ApJ*, **689**, L69  
 Spruit, H. C. 1981, *A&A*, **98**, 155  
 Stangalini, M. 2014, *A&A*, **561**, L6  
 Stangalini, M., Berrilli, F., & Consolini, G. 2013a, *A&A*, **559**, A88  
 Stangalini, M., Solanki, S. K., Cameron, R., & Martínez Pillet, V. 2013b, *A&A*, **554**, A115  
 Stangalini, M., Consolini, G., Berrilli, F., De Michelis, P., & Tozzi, R. 2014, *A&A*, **569**, A102  
 Steiner, O., Grossmann-Doerth, U., Knoelker, M., & Schüssler, M. 1998, *ApJ*, **495**, 468  
 van Noort, M., Rouppe van der Voort, L., & Löfdahl, M. G. 2005, *Sol. Phys.*, **228**, 191  
 Vigeesh, G., Fedun, V., Hasan, S. S., & Erdélyi, R. 2012, *ApJ*, **755**, 18  
 Welsch, B. T., & Longcope, D. W. 2003, *ApJ*, **588**, 620