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PTF10nvg: AN OUTBURSTING CLASS I PROTOSTAR IN THE PELICAN/NORTH AMERICAN NEBULA

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ABSTRACT

During a synoptic survey of the North American Nebula region, the Palomar Transient Factory (PTF) detected an optical outburst (dubbed PTF10nvg) associated with the previously unstudied flat or rising spectrum infrared source IRAS 20496+4354. The PTF *R*-band light curve reveals that PTF10nvg brightened by more than 5 mag during the current outburst, rising to a peak magnitude of $R_{\text{PTF}} \approx 13.5$ in 2010 September. Follow-up observations indicate that PTF10nvg has undergone a similar ~ 5 mag brightening in the *K* band and possesses a rich emission-line spectrum, including numerous lines commonly assumed to trace mass accretion and outflows. Many of these lines are blueshifted by ~ 175 km s⁻¹ from the North American Nebula's rest velocity, suggesting that PTF10nvg is driving an outflow. Optical spectra of PTF10nvg show several TiO/VO band heads fully in emission, indicating the presence of an unusual amount of dense ($>10^{10}$ cm⁻³), warm (1500–4000 K) circumstellar material. Near-infrared spectra of PTF10nvg appear quite similar to a spectrum of McNeil's Nebula/V1647 Ori, a young star which has undergone several brightenings in recent decades, and 06297+1021W, a Class I protostar with a similarly reached near-infrared emission line spectrum. While further monitoring is required to fully understand this event, we conclude that the brightening of PTF10nvg is indicative of enhanced accretion and outflow in this Class-I-type protostellar object, similar to the behavior of V1647 Ori in 2004–2005.

Key words: stars: emission-line, Be – stars: formation – stars: pre-main sequence – stars: variables: general – stars: winds, outflows

Online-only material: color figures

1. INTRODUCTION

A defining characteristic (e.g., Joy 1945) of young stars is their photometric variability. At optical wavelengths, periodic low-amplitude variations are thought to arise from surface inhomogeneities on a rotating star; aperiodic large-amplitude variability is more often attributed to accretion-related activity. The “activity” amplitude generally declines as the star ages and the disk evolves, with “outburst” behavior exhibited by some of the most active sources. This has resulted in the definition of EXor type (actually named after EX Lup) young-star variables; these objects often undergo repeated outbursts with 2–3 mag amplitudes and characteristic timescales of weeks to months (Herbig et al. 2001; Herbig 2008; Lorenzetti et al. 2009).

More extreme are the FU Ori-type variables, which undergo 4–6 mag amplitude increases over a few months, followed by a slow, decades-long decay (Herbig 1989; Hartmann & Kenyon 1996). In both cases, the variability is commonly attributed to nonsteady mass accretion from a circumstellar disk. Theoretical models typically invoke thermal, gravitational, or magnetorotational disk instabilities to explain these accretion outbursts (Bell & Lin 1994; Kley & Lin 1999; Armitage et al. 2001; Vorobyov & Basu 2005; Boley et al. 2006; Zhu et al. 2009), potentially triggered by binary interactions (Reipurth & Aspin 2004b; Lodato & Clarke 2004). Recently documented outbursting young stars include V1647 Ori (McNeil et al. 2004; Briceño et al. 2004; Reipurth & Aspin 2004a; Walter et al. 2004; Ojha et al. 2006; Fedele et al. 2007; Aspin et al. 2009), V1118 Ori (Lorenzetti et al. 2007; Audard et al. 2010), V733 Cep (Reipurth et al. 2007; Peneva et al. 2010), ZCMa (Szeifert et al. 2010; Whelan et al. 2010), and EX Lup itself (Herbig et al. 2001).

With optical and near-infrared telescopes now regularly patrolling the sky, and ongoing variability surveys that target active star-forming regions, the full gamut of young-star variability will soon be revealed, quantifying the characteristic timescales

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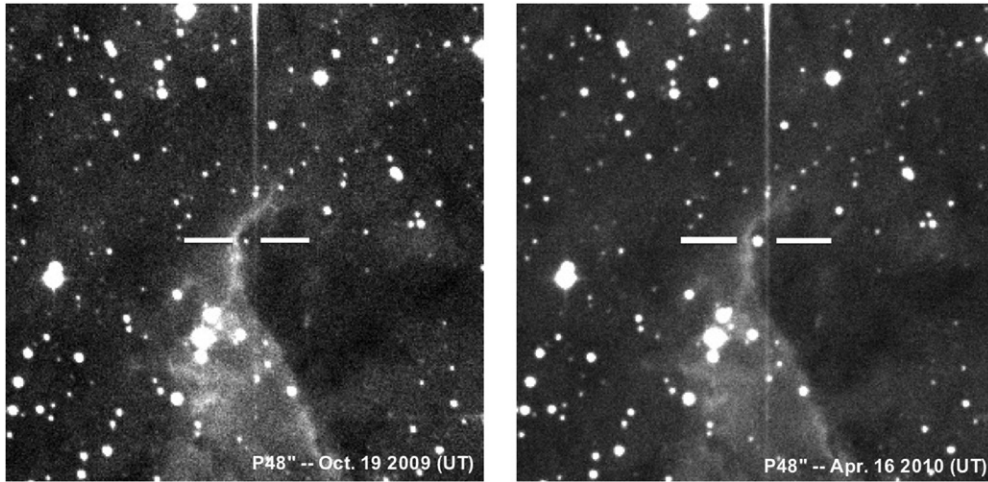


Figure 1. PTF images of PTF10ngv on 2009 October 19 (left) and nearly 6 months later on 2010 April 16 (right). Images are $5'$ on a side, with north up and east to the left; PTF10ngv is identified with white crosshairs. The nebulosity in each R -band frame was pre-existing and is not associated with the outbursting behavior of the point source. PTF10ngv is embedded within the molecular cloud associated with the North American/Pelican Nebula. A bright diffraction spike from a star just north of PTF10ngv does not substantially affect the PSF photometry derived for PTF10ngv.

for eruptive events of varying amplitudes. One such ongoing investigation is using the Palomar 48 inch Schmidt telescope and the Palomar Transient Factory (PTF; Law et al. 2009; Rau et al. 2009) infrastructure to catalogue photometric variability in the North American/Pelican Nebula region of recent star formation. Located at a distance of ~ 600 pc, the population of young stars in this region (identified largely from surveys in $H\alpha$; e.g., Herbig 1958; Welin 1973; Ogura et al. 2002, or the mid-IR, e.g., Guieu et al. 2009; L. M. Rebull et al. 2011, in preparation) is not as well studied as members of more proximate star-forming regions, but already includes a disproportionately large number of the known members of the exclusive FU Ori class, namely V1057 Cyg and V1515 Cyg. These two objects, along with FU Ori itself, are the defining members of the class; they exhibit strong wind signatures in lines such as Na I D and $H\alpha$, metallic absorption patterns similar to those of FG-type supergiants in the optical and M-type supergiants in the IR, and large thermal IR luminosity.

We have identified a new outbursting object located at position $\alpha = 20^{\text{h}}51^{\text{m}}26^{\text{s}}.23$, $\delta = +44^{\circ}05'23''.9$ (J2000.0), just south of the Pelican Nebula. It sits within the apex of a nebular arc that wraps around the object from northwest to the east and then to the south. The source is located within $1''$ of the position of IRAS 20496+4354, well inside the $16''$ error ellipse associated with that detection. The source was further detected by the Midcourse Space Experiment (MSX) in 1996–1997, *Spitzer* in 2006 August 9–11 (Guieu et al. 2009; L. M. Rebull et al. 2011, in preparation), and AKARI/IRC during its 2006–2007 mission.

At shorter wavelengths, the source was not visible ($K > 15.3$ mag) to the Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006). The 2–24 μm spectral index ($\alpha = d \log(\lambda F_\lambda) / d \log \lambda$; Lada 1987) implied by the 2MASS K -band upper limit and the IRAS 24 μm detection is $\alpha \gtrsim 1.7$, consistent with a designation as a heavily embedded Class I protostar in the classification system devised by Lada (1987). The source does, however, appear in the J and K images of Magnier et al. (1999) that were obtained as follow-up observations to the IRAS detection. The source is not obvious on Digital Sky Survey (DSS) images (implying $m_{\text{pg}} > 21$ mag), but NOMAD, using DSS-2 data, reports $m_B = 20.17$ and $m_R = 18.28$ mag for a source $3''.4$ away from the source position reported here (Zacharias et al. 2005). Images of the field taken with $H\alpha$ and

[S II] filters by Bally & Reipurth (2003) appear to show faint smudges at the source position.

We report here the outburst discovery data, as well as follow-up multi-filter optical and near-infrared (NIR) photometric and spectroscopic monitoring that was obtained during various stages of the 2010 outburst. We find evidence for strong emission lines and broadly absorbed wind features, similar to other outbursting young objects, but also some unique spectroscopic attributes, including molecular emission from TiO and VO. We analyze the data in the context of an eruptive Class-I type, heavily embedded protostar.

2. OBSERVATIONS

2.1. Optical Photometry

During the 2009 and 2010 observing seasons, PTF obtained red optical images of the North American/Pelican Nebula star-forming region with a typical five-day cadence. These observations were conducted with the main PTF Survey Camera, the former CFHT12K mosaic camera now extensively re-engineered and mounted on the 48 inch Samuel Oschin telescope at Palomar Observatory (hereafter “P48”). The camera is a mosaic of 11 CCDs, covering a 7.8 deg^2 field of view with $1''$ sampling; typical conditions at Palomar Observatory produce $2''.0$ full width at half-maximum (FWHM) intensity images (Law et al. 2009). The R_{PTF} filter, a Mould R band, is similar to SDSS- r in shape but is shifted redward by $\sim 27 \text{ \AA}$ and is $\sim 20 \text{ \AA}$ wider. The typical 5σ limiting magnitudes are $m_R \approx 20.6$ (AB) in 60 s exposures. Representative images of PTF10ngv from the 2009 and 2010 observing seasons are shown in Figure 1.

Transient sources are detected in PTF monitoring data by means of automated reduction pipelines, including a near-real-time image-subtraction pipeline hosted at Lawrence Berkeley National Laboratory (LBNL). Well-detected sources in the difference images are scored (using a human trained machine-based classifier) for their likelihood of being truly astrophysical in nature or of spurious origin. Variable sources with larger likelihoods of being nonspurious are passed to an automatic source classifier at UC Berkeley (“Oarical”), which combines PTF measurements with all other available information (e.g., SIMBAD identifications, 2MASS photometry, etc.) to provide probabilistic classifications of PTF detections (J. S. Bloom

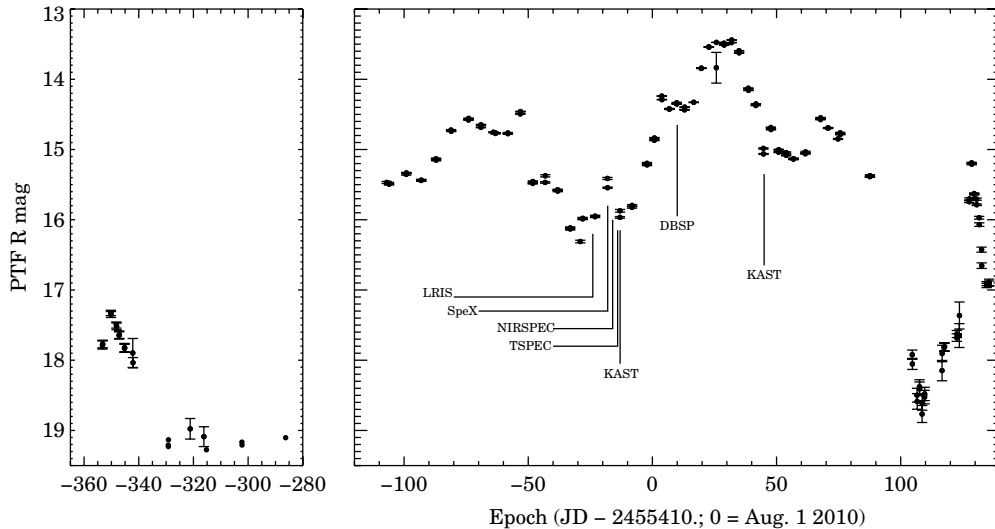


Figure 2. R_{PTF} -band light curve of PTF10nvg, measured via PSF photometry of individual P48 frames (see Section 2.1 for details). Left panel shows observations from the 2009 season; right panel shows 2010 observations, with labels identifying the timing of the spectra presented in Sections 2.3 and 2.4

et al. 2011, in preparation). These initial classifications are made available to PTF collaboration members via the PTF Follow-up Marshal, which enables visual inspection of current and reference images, precursor PTF light curves, and any subsequent PTF spectroscopy.

The source reported here was detected by the PTF pipeline and automatically assigned the name 10nvg, following PTF naming conventions. PTF images are typically re-reduced for sources of particular interest, such as PTF10nvg, via a “white-glove pipeline,” where PSF photometry is performed on individual frames. The PSF photometry is calculated using a modified version of the pipeline developed by the Supernova Legacy Survey (see Astier et al. 2006). Absolute calibration is performed relative to the USNO-B1 catalog (Monet et al. 2003), and typically has an uncertainty of ~ 0.15 mag. The full P48 light curve of PTF10nvg produced by this “white-glove reduction” is shown in Figure 2 and the photometry appears in Table 1.

The PTF also makes use of the robotic Palomar 60 inch telescope (P60; Cenko et al. 2006) to obtain multi-color photometry for source verification and classification purposes. PTF10nvg was observed by the robotic Palomar 60-inch starting 2010 August 1.35 in the r -band, i -band, and z -band filters. The absolute zero-point calibration was derived relative to the Sloan Digital Sky Survey (SDSS) from separate fields observed by the Palomar 60 inch during the night of 2010 September 3 with the same filter. All other fields are calibrated relative to this night. Figure 3 presents a three-color image constructed from the P60 riz frames, and Table 2 contains the individual measurements.

We note that there is a ~ 0.45 mag offset between the P48 R_{PTF} and P60 r -band light curves for PTF10nvg. A simple zero-point error may contribute to this offset, but it is likely dominated by differences in each telescope’s filter + detector response, and the different photometric systems underlying their calibration (Vega R versus AB magnitude r). Given the multiple possible contributions to this offset, we have chosen to preserve each light curve on its native system, rather than explicitly transform them onto a common system.

2.2. Near-infrared Photometry

Near-infrared observations of PTF10nvg were conducted with the 1.3 m Peters Automated Infrared Imaging Telescope (PAIRITEL; Bloom et al. 2006) starting on 2010 July 10.

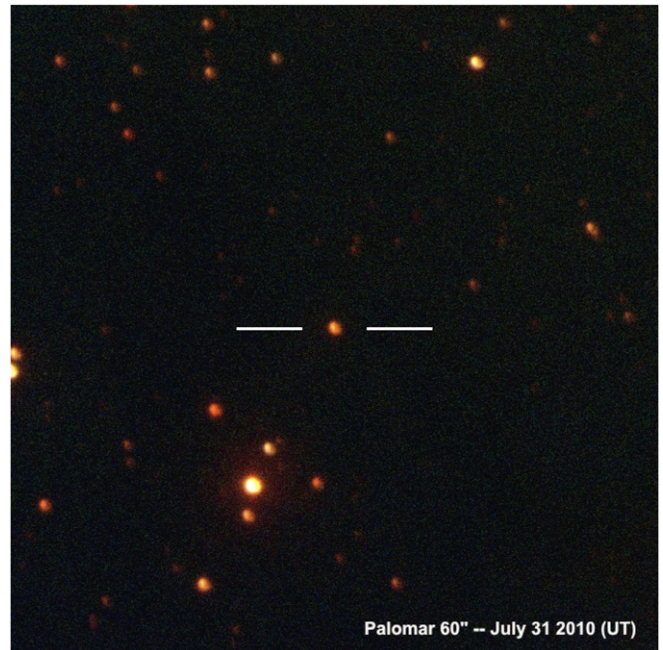


Figure 3. Three-color frame of PTF10nvg riz images obtained with the P60. The frame is 4:1 on a side, with north up and east to the left. (A color version of this figure is available in the online journal.)

PAIRITEL is a roboticized system using the former 2MASS southern hemisphere survey camera that employs two dichroics to observe simultaneously in the J , H , and K_s bands. Observations were scheduled and executed via a robotic system. PAIRITEL is operated in a fixed observing mode in which 7.8 s double-correlated images are created from the difference of a 7.851 s and a 51 ms integration taken in rapid succession (see Blake et al. 2008). The standard observing procedure involves taking three image pairs prior to dithering the telescope.

The raw data from these images are reduced using standard IR reduction methods via PAIRITEL PIPELINE III and the flux for all sources is measured via aperture photometry using SExtractor (Bertin & Arnouts 1996), calibrated against 2MASS. PTF10nvg saturates the 7.851 s frames; however, PIPELINE III produces “short-frame” mosaics consisting of reduced, stacked

Table 1
P48 R-Band Photometry of PTF10nvg

Epoch (JD -2,400,000)	mag	1 σ Error
55056.7205	17.784	0.058
55056.7954	17.768	0.054
55059.7615	17.327	0.038
55059.8570	17.349	0.042
55061.8156	17.499	0.043
55061.8834	17.519	0.048
55062.8266	17.648	0.051
55062.8906	17.637	0.053
55064.7860	17.831	0.058
55064.8311	17.820	0.059
55067.8178	17.897	0.206
55067.8929	18.035	0.073
55080.7902	19.209	99.000
55080.8339	19.229	99.000
55080.8608	19.133	99.000
55088.8156	18.976	0.147
55093.8177	19.088	0.141
55094.7525	19.275	99.000
55107.7216	19.166	99.000
55107.7656	19.207	99.000
55123.7247	19.103	99.000
55302.9792	15.469	0.014
55303.9110	15.481	0.013
55303.9593	15.494	0.011
55310.8700	15.354	0.012
55310.9130	15.334	0.008
55316.8754	15.441	0.013
55316.9198	15.437	0.012
55322.8590	15.132	0.011
55322.9028	15.151	0.010
55328.9085	14.723	0.007
55328.9530	14.739	0.006
55335.9332	14.579	0.006
55335.9774	14.557	0.006
55340.9292	14.685	0.006
55340.9735	14.648	0.007
55345.9696	14.754	0.008
55346.7699	14.765	0.007
55346.8139	14.771	0.007
55351.7712	14.774	0.008
55351.8145	14.764	0.007
55356.8089	14.493	0.005
55356.8533	14.460	0.005
55361.8110	15.482	0.009
55361.8546	15.456	0.010
55366.8177	15.469	0.011
55366.8615	15.374	0.015
55371.8399	15.573	0.009
55371.8836	15.592	0.009
55376.8771	16.131	0.016
55376.9201	16.117	0.015
55380.9668	16.310	0.019
55381.9049	15.987	0.013
55381.9551	15.980	0.013
55386.8972	15.956	0.013
55386.9442	15.949	0.013
55391.8996	15.412	0.015
55391.9437	15.544	0.009
55396.9115	15.966	0.012
55396.9655	15.873	0.019
55401.8040	15.824	0.011
55401.8483	15.796	0.012
55407.7945	15.197	0.013
55407.8421	15.218	0.012
55410.7952	14.868	0.005

Table 1
(Continued)

Epoch (JD - 2,400,000)	mag	1 σ Error
55410.8404	14.836	0.008
55413.7918	14.239	0.004
55413.8356	14.290	0.004
55416.7871	14.420	0.005
55416.8303	14.425	0.005
55419.7818	14.352	0.004
55419.8261	14.334	0.005
55422.9468	14.435	0.005
55422.9910	14.396	0.005
55426.6605	14.327	0.005
55429.6767	13.845	0.002
55429.7202	13.842	0.003
55432.7443	13.539	0.002
55432.7874	13.542	0.002
55435.7587	13.836	0.219
55435.8028	13.475	0.002
55438.8303	13.487	0.002
55438.9133	13.512	0.002
55441.8654	13.478	0.002
55441.9300	13.441	0.002
55444.8899	13.594	0.002
55444.9340	13.623	0.002
55448.6442	14.128	0.004
55448.6889	14.157	0.004
55451.6560	14.372	0.004
55451.7002	14.351	0.004
55454.7211	14.985	0.010
55454.7651	15.063	0.007
55457.7168	14.687	0.007
55457.7608	14.715	0.006
55460.8389	15.037	0.007
55460.8875	15.002	0.007
55463.8407	15.043	0.010
55463.8859	15.083	0.008
55466.8255	15.137	0.008
55466.8771	15.132	0.009
55471.6094	15.060	0.007
55471.6884	15.035	0.010
55477.6909	14.573	0.004
55477.7342	14.549	0.006
55480.7401	14.696	0.007
55480.7835	14.692	0.006
55484.7627	14.849	0.008
55485.6193	14.763	0.007
55485.6635	14.784	0.006
55497.5978	15.386	0.010
55497.6412	15.372	0.013
55514.6697	17.921	0.066
55514.7179	18.053	0.079
55516.6017	18.498	0.098
55516.6451	18.587	0.111
55517.5770	18.378	0.100
55517.6203	18.407	0.097
55518.6247	18.766	0.121
55518.6680	18.607	0.105
55519.6221	18.525	0.097
55519.6656	18.480	0.095
55526.6565	18.148	0.145
55526.7003	17.904	0.117
55527.6085	17.810	0.056
55527.6530	17.813	0.058
55532.5921	17.678	0.053
55532.6356	17.627	0.050
55533.6257	17.649	0.170
55533.6690	17.363	0.192

Table 1
(Continued)

Epoch (JD – 2,400,000)	mag	1 σ Error
55537.6163	15.743	0.014
55537.6606	15.702	0.014
55538.5897	15.190	0.009
55538.6337	15.208	0.008
55539.5941	15.627	0.009
55539.6384	15.638	0.010
55540.6025	15.707	0.015
55540.6464	15.786	0.011
55541.6075	15.971	0.019
55541.6511	16.069	0.025
55542.5817	16.426	0.036
55542.6250	16.654	0.040
55544.5990	16.909	0.028
55544.6459	16.935	0.030
55545.6112	16.892	0.046
55545.6554	16.907	0.035

51 ms images (see also Bloom et al. 2009). The “short-frame” mosaics contain >10 bright 2MASS stars which we use to properly calibrate photometric measurements of PTF10nvg in these images. PAIRITEL has a known systematic uncertainty of ~ 0.02 – 0.03 mag in each of the J , H , and K_s bands (see also Blake et al. 2008; Perley et al. 2010), which, in the case of PTF10nvg, is larger than the statistical error in all three bands. Thus, we add a systematic error of 0.03 mag in quadrature with the statistical uncertainty to determine the total uncertainty in each band. The JHK_s images of PTF10nvg obtained by PAIRITEL on 2010 July 10 are shown as a three-color frame in Figure 4; all JHK_s magnitudes measured for PTF10nvg by PAIRITEL are presented in Table 3.

We also examined the 2MASS data products to determine if PTF10nvg was detected during that survey. A query of the 2MASS source catalog via the GATOR interface does not identify a source within $10''$ of this position, and visual inspection of the 2MASS images confirm that the source was undetected in 2000 October, when the field was imaged. Analysis of the 2MASS frames provides 5σ upper limits on the brightness of the source at the epoch of the 2MASS observations of $J > 16.35$, $H > 15.38$, and $K_s > 14.80$; we include these upper limits in the JHK_s photometry presented in Table 3.

2.3. Optical Spectroscopy

Moderate-resolution optical spectra of PTF10nvg were obtained on four occasions: 2010 July 8 with the Low Resolution Imaging Spectrometer (LRIS; Oke et al. 1995; McCarthy et al. 1998; Steidel et al. 2004) on the Keck-I 10 m telescope, 2010 July 19 and 2010 September 16 with the Kast double spectrograph (Miller & Stone 1993) on the Shane 3 m telescope at Lick Observatory, and 2010 August 11 with the Double Spectrograph (DBSP; Oke & Gunn 1982) on the Hale 5 m telescope at Palomar Observatory.

LRIS observations were made by J.S.B., A. Cucchiara, A. N. Morgan, and D. A. Perley with the D560 dichroic, which has a 50% transmission split at $\lambda \approx 5700$ Å, and a $1''$ long slit. The blue side of the spectrograph was configured with the 600/4000 grism, providing ~ 4 Å resolution over the 3010–5600 Å wavelength range. The red spectrograph utilized the 400/8500 grating, providing ~ 7 Å resolution over

Table 2
P60 Optical Photometry of PTF10nvg

Epoch (JD – 2,400,000)	Filter Filter	mag mag	1 σ Error
55409.353	<i>r</i>	15.487	0.170
55409.355	<i>z</i>	13.274	0.254
55410.177	<i>z</i>	13.134	0.332
55411.234	<i>z</i>	13.096	0.189
55412.206	<i>i</i>	14.208	0.025
55412.207	<i>z</i>	12.830	0.225
55413.326	<i>r</i>	14.776	0.150
55413.327	<i>z</i>	12.772	0.258
55414.255	<i>z</i>	12.864	0.266
55415.280	<i>z</i>	12.908	0.230
55416.279	<i>i</i>	14.218	0.043
55416.280	<i>z</i>	12.935	0.196
55417.244	<i>r</i>	14.920	0.159
55417.245	<i>z</i>	12.859	0.285
55418.377	<i>z</i>	12.947	0.240
55419.316	<i>z</i>	12.856	0.215
55420.184	<i>i</i>	14.105	0.040
55420.185	<i>z</i>	12.795	0.238
55421.267	<i>r</i>	15.059	0.162
55421.268	<i>z</i>	13.026	0.267
55423.183	<i>z</i>	12.994	0.188
55424.408	<i>i</i>	14.241	0.034
55424.410	<i>z</i>	12.896	0.257
55433.350	<i>i</i>	13.360	0.113
55433.351	<i>r</i>	13.877	0.191
55433.352	<i>z</i>	12.661	0.238
55434.286	<i>z</i>	11.860	0.546
55436.256	<i>z</i>	12.001	0.226
55438.220	<i>i</i>	13.285	0.226
55438.298	<i>r</i>	13.969	0.121
55438.299	<i>z</i>	12.089	0.153
55442.213	<i>i</i>	13.246	0.002
55442.214	<i>r</i>	13.899	0.002
55442.214	<i>z</i>	12.598	0.001
55443.218	<i>i</i>	13.219	0.018
55443.220	<i>r</i>	13.798	0.160
55443.225	<i>z</i>	11.963	0.250
55444.239	<i>i</i>	13.402	0.029
55444.240	<i>r</i>	13.978	0.186
55444.241	<i>z</i>	12.172	0.233
55445.261	<i>i</i>	13.572	0.020
55445.265	<i>r</i>	14.176	0.181
55445.277	<i>z</i>	12.336	0.231
55446.280	<i>i</i>	13.587	0.020
55446.281	<i>r</i>	14.194	0.181
55446.282	<i>z</i>	12.311	0.230
55447.236	<i>i</i>	13.875	0.029
55447.237	<i>z</i>	12.576	0.219
55448.164	<i>z</i>	12.693	0.196
55449.372	<i>z</i>	12.746	0.137
55450.161	<i>r</i>	14.513	0.161
55450.162	<i>z</i>	12.537	0.282
55451.166	<i>r</i>	14.863	0.129
55452.355	<i>z</i>	13.002	0.225
55458.340	<i>r</i>	15.241	0.159
55458.341	<i>z</i>	13.158	0.309
55459.311	<i>z</i>	13.334	0.200
55460.228	<i>z</i>	13.455	0.079
55461.197	<i>z</i>	13.382	0.144
55462.270	<i>r</i>	15.534	0.181
55462.271	<i>z</i>	13.463	0.122
55463.187	<i>z</i>	13.463	0.116
55469.183	<i>r</i>	15.763	0.190
55469.184	<i>z</i>	13.611	0.179
55470.362	<i>z</i>	13.576	0.160

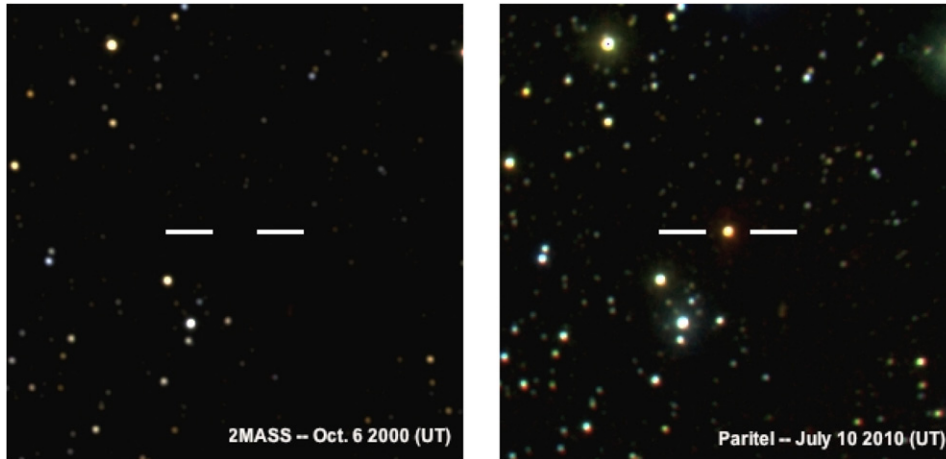


Figure 4. Three-color *JHK* images of the field containing PTF10nvg obtained by 2MASS (left) and PAIRITEL (right). Frames are 5' on a side, with north up and east to the left. PTF10nvg is not detected in the 2MASS imaging acquired on 2000 October 6 but is visible in the PAIRITEL imaging acquired on 2010 July 10.

(A color version of this figure is available in the online journal.)

Table 2
(Continued)

Epoch (JD – 2,400,000)	Filter	mag	1 σ Error
55473.308	<i>r</i>	15.486	0.139
55473.309	<i>z</i>	13.421	0.138
55477.322	<i>r</i>	15.059	0.161
55477.324	<i>z</i>	13.042	0.166
55478.146	<i>z</i>	13.129	0.148
55479.117	<i>z</i>	13.047	0.075
55480.171	<i>z</i>	12.585	0.547
55481.147	<i>r</i>	15.335	0.089
55481.148	<i>z</i>	13.250	0.099
55482.108	<i>z</i>	13.317	0.160
55483.128	<i>z</i>	13.356	0.080
55484.182	<i>z</i>	13.360	0.179
55485.109	<i>r</i>	15.283	0.146
55485.109	<i>z</i>	13.321	0.177
55496.280	<i>r</i>	16.102	0.233
55496.281	<i>z</i>	14.078	0.228
55497.130	<i>z</i>	13.919	0.174
55499.122	<i>z</i>	13.969	0.103

Table 3
NIR Photometry of PTF10nvg

Epoch (JD)	<i>J</i> mag	<i>J</i> Error	<i>H</i> mag	<i>H</i> Error	<i>K</i> mag	<i>K</i> Error
2451823.66652	>16.35	...	>15.38	...	>14.80	...
2455387.87633	12.22	0.07	10.26	0.02	8.45	0.01
2455392.79441	12.23	0.01	10.38	0.06	8.45	0.02
2455450.69076	10.88	0.02	9.39	0.01	8.06	0.01
2455463.65116	11.48	0.03	9.86	0.01	8.39	0.01
2455466.62568	11.62	0.03	10.02	0.01	8.38	0.01

5600–10300 Å. The LRIS spectra were reduced in the IRAF¹⁵ environment using standard routines. Cosmic rays were removed using the L.A. Cosmic routine (van Dokkum 2001). Spectra were extracted optimally (Horne 1986) and wavelength calibration was performed first relative to arc lamps and then tweaked based on night-sky lines in each individual image. Both air-to-vacuum and heliocentric corrections were then ap-

¹⁵ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

plied to all spectra. Extracted spectra were divided through by a smoothed flux standard to remove narrow-band (<50 Å) instrumental effects (Bessell 1999). Finally, telluric absorption features were removed using the continuum from spectrophotometric standards (Wade & Horne 1988; Matheson et al. 2000).

The July 19 Kast observations were obtained by C. V. Griffith and M. T. Kandrashoff, with reductions performed by J.M.S.; the September 16 observations were obtained by S.B.C. and M. T. Kandrashoff, with reductions performed by S.B.C. Kast was configured with a 2'' slit, a 600/4310 grism on the blue side, and a 300/7500 grating on the red side, yielding resolutions of ~4 and ~10 Å, respectively.

DBSP data were acquired by L.M.R. using the D55 dichroic to obtain spectra with a 600 line mm⁻¹ grating blazed at 4000 Å from the atmospheric cutoff to 5500 Å at 1.1 Å resolution, and with a 158 line mm⁻¹ grating blazed at 7500 Å over the range 6300–8800 Å at 2.5 Å dispersion. DBSP data reduction was performed by L.A.H., including flat fielding, extraction, and wavelength calibration, using standard routines in the IRAF environment.

2.4. Near-infrared Spectroscopy

Moderate-resolution NIR spectra of PTF10nvg were obtained on 2010 July 14 with the SpeX spectrograph (Rayner et al. 2003) at the NASA Infrared Telescope Facility, and on 2010 July 18 with the TripleSpec Spectrograph (Wilson et al. 2004) at the Apache Point Observatory.

SpeX observations were obtained by J.R. on 2010 July 14 with a 0'3 slit under clear skies and good seeing (0'4–0'5). Both the short-cross-dispersed (SXD) and long cross-dispersed (LXD) modes were utilized, providing nearly contiguous coverage over the range 0.8–2.5 μm with a spectral resolution of $R = 2000$, and over 3–4.2 μm (*L* band) and 4.5–5.0 μm (*M* band) at $R \approx 2500$. PTF10nvg was nodded along the slit during the observation, such that differencing consecutive images effectively removes the sky emission. The spectra were differenced, flat fielded, extracted, and wavelength calibrated using the SpeXTool IDL routines (Cushing et al. 2004). Telluric correction and absolute flux calibration were performed using a spectrum of a nearby A0V star obtained at similar air mass and time to PTF10nvg and reduced in similar fashion; they were applied using the XTELLCOR IDL package (Vacca et al. 2003).

Table 4
Pre-outburst Photometry of PTF10ngv

Wavelength (μm)	Flux Density (Jy)	Flux Error (1σ)	Observatory	Epoch	Reference
12	3.39	0.33	IRAS	1983	Odenwald (1989)
25	6.59	0.66	IRAS	1983	Odenwald (1989)
60	27.89	1.0	IRAS	1983	Odenwald (1989)
100	57.35	1.0	IRAS	1983	Odenwald (1989)
8.28	1.6	0.66	MSX	1996–1997	Egan et al. (1999)
12.13	2.49	0.137	MSX	1996–1997	Egan et al. (1999)
14.65	3.045	0.189	MSX	1996–1997	Egan et al. (1999)
21.34	3.14	0.201	MSX	1996–1997	Egan et al. (1999)
3.6	0.083820	0.0041974110	Spitzer/IRAC	2006–2007	Guieu et al. (2009)
4.5	0.231500	0.0011582866	Spitzer/IRAC	2006–2007	Guieu et al. (2009)
5.8	0.609100	0.0030475698	Spitzer/IRAC	2006–2007	Guieu et al. (2009)
8.0	1.171000	0.0058589830	Spitzer/IRAC	2006–2007	Guieu et al. (2009)
24.0	2.222000	0.088946621	Spitzer/MIPS	2006–2007	L. M. Rebull et al. 2011 (in preparation)
70.0	6.732000	...	Spitzer/MIPS	2006–2007	L. M. Rebull et al. 2011 (in preparation)
9.0	1.96	0.2	AKARI	2006–2007	Ishihara et al. (2010)
15.0	3.875	0.457	AKARI	2006–2007	Ishihara et al. (2010)

TripleSpec observations were obtained by K.R.C. on 2010 July 18 with the $0''.7$ slit, providing nearly contiguous coverage over $1.0\text{--}2.5\ \mu\text{m}$ at $R \approx 5000$. Spectra were obtained and reduced in the same manner as the SpeX observations, using a version of the SpeXTool package modified for use with APO/TripleSpec data.

A NIRSPEC spectrum covering $0.95\text{--}1.12\ \mu\text{m}$ at $R \approx 25,000$ was obtained on 2010 July 16 by L.A.H. Standard data-reduction steps including wavelength calibration and image rectification, extraction of one-dimensional spectra, and removal of telluric emission and absorption features were carried out by W.J.F. within the REDSPEC IDL package developed by S. S. Kim, L. Prato, and I. McLean.

3. DISCOVERY AND CHARACTERIZATION

PTF10ngv was first identified as an optical transient by the PTF automatic discovery and classification codes on 2010 July 8. It was selected for spectroscopic follow-up observation as a bright event observable with Keck during morning twilight of July 8, and subsequently visually identified as a likely outbursting young star. Analysis of the PTF data and the follow-up photometry and spectroscopy is presented below.

3.1. Photometric Analysis

3.1.1. Outburst Characterization from the Optical/PTF Light Curve

Although PTF10ngv is not known historically as an optical source, it was detected by PTF at $m_R \approx 18$ mag in mid-2009 but then faded to ~ 19.25 mag on a timescale of roughly 2 months. It brightened again by more than 4 mag in less than 200 days, while the field was too close to the Sun for observations. PTF10ngv reached $m_R = 15$ early in 2010, faded to 16.3 mag within 40 days of the first peak, and then brightened again to 13.5 mag before fading to ~ 15.5 mag in early 2010 November. The two brightness peaks, as well as the intermediate lull, are all ~ 50 days in width.

3.1.2. Pre- and Post-outburst SED

Pre-outburst photometry that we have collated from the literature is presented in Table 4, and the spectral energy distributions (SEDs) assembled from this pre-outburst photometry and our own post-outburst detection data are displayed in Figure 5.

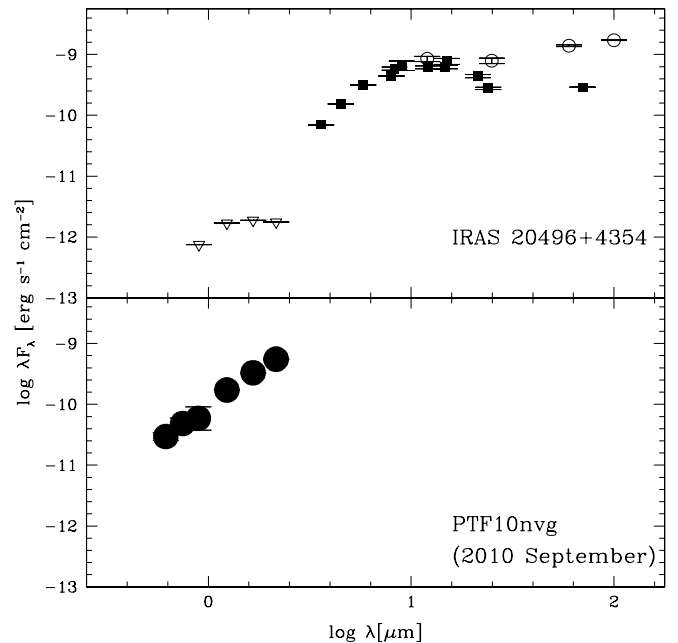


Figure 5. Pre-outburst (top panel) and post-outburst (bottom panel) SEDs of PTF10ngv. The pre-outburst data derive from KPNO 0.9 m optical and 2MASS NIR upper limits (triangles), *Spitzer* IRAC (3.6, 4.5, 5.8, 8.0 μm ; filled squares) and MIPS (24, 70 μm ; filled squares), MSX (8.3, 12.1, 14.6, 21.3 μm ; filled squares), AKARI (9 and 18 μm ; filled squares), and *IRAS* (12, 25, 60, 100 μm ; open circles) photometric measurements (note that the *IRAS* beam was larger than that of the other instruments). The post-outburst data are those of 2010 July 9 as reported in Tables 1 and 3, when the object was in a relative lull between its two optical brightness maxima (see Figure 2).

From 3 to 10 μm , the SED is rising then becomes roughly flat from 10 to 100 μm . The integrated luminosity of the pre-outburst IR source between 3 and 100 μm is $25 L_{\odot}$; extrapolating the SED longward of the *IRAS* 100 μm measurement as a standard Rayleigh–Jeans function, and integrating to $\lambda = \infty$, increases the luminosity by $\sim 40\%$. Other Class I sources, however, have SEDs that peak at or slightly longward of 100 μm (e.g., Enoch et al. 2009); if PTF10ngv’s SED possesses a similar peak, the total luminosity would be larger than this, which assumes a sharply declining power-law directly from 100 μm . The integrated luminosity of the outbursting source between 0.7 and

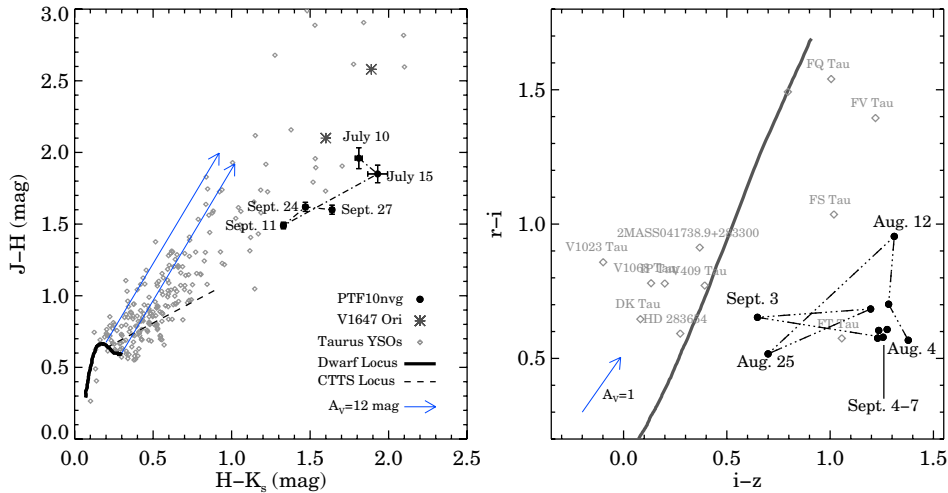


Figure 6. Location of PTF10ngv in JHK_s (left) and riz (right) color-color diagrams. Solid lines in each panel indicate the locus of SDSS/2MASS colors for main-sequence field dwarfs (F to M spectral types; Covey et al. 2007); the dashed line in the top panel shows the Classical T Tauri locus derived by Meyer et al. (1997) and transformed into the 2MASS photometric system using the relationship derived by Carpenter (2001). Class I, II, and III objects in Taurus, as well as the outbursting object V1647 Ori, are shown for comparison; photometry for these comparison stars was catalogued by Finkbeiner et al. (2004), Reipurth & Aspin (2004a), and Rebull et al. (2010).

(A color version of this figure is available in the online journal.)

$2.2 \mu\text{m}$ is $2.5 L_{\odot}$, compared to $<0.016 L_{\odot}$ before the outburst. There is no post-outburst photometry longward of $2.2 \mu\text{m}$.

The optical and NIR colors of PTF10ngv, shown in Figure 6, are extremely red during outburst: $r - i \approx 0.6 \text{ mag}$, $i - z \approx 1.1 \text{ mag}$, $J - H \approx 1.9 \text{ mag}$, and $H - K_s \approx 1.9 \text{ mag}$. These colors are within the distribution but toward the blue end of NIR colors exhibited by Class-I-type stars in Taurus, and significantly redder than can be explained by a range of disk accretion, inner hole, and inclination properties along the Class II “Classical T Tauri locus” (Meyer et al. 1997). In 2010 July, PTF10ngv appeared redder in $H - K_s$ and bluer in $J - H$ than most known “FU Ori-like” objects (e.g., see Figure 4 by Greene et al. 2008). As PTF10ngv began fading in 2010 September, however, its observed JHK_s colors became somewhat bluer, comparable to the colors of V346 Nor, a previously known “FU Ori-like” object. The JHK_s colors of PTF10ngv are also similar to those of V1647 in 2004 February–March, early in the 2004–2005 outburst.

High-energy radiation has been detected from other outbursting young stars (e.g., V1647; Kastner et al. 2006), but no X-ray detection can be confidently associated with PTF10ngv at this time. X-ray observations of PTF10ngv were obtained with the *Swift* X-Ray Telescope (XRT; Burrows et al. 2005) starting on 2010 August 29 for a total exposure time of 3.6 ks. We extract the 0.5–8.0 keV counts from an extraction region of 64 pixels, $\sim 2.5 \text{ arcmin}$, where we fit the point spread function model (see Butler & Kocevski 2007) at the centroid of the optical position of PTF10ngv. We detect no X-ray emission from PTF10ngv. Assuming a $\Gamma = 2$ power-law spectrum and a column density of $N_{\text{H}} = 10^{22} \text{ cm}^{-2}$ we find a 3σ upper limit of the flux: $F_X < 1.2 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$. Consistent with both the sensitivity of the *ROSAT* All Sky Survey (Voges et al. 1999) and the *Swift* upper limit, the closest *ROSAT* source is approximately $0'9 \text{ SW}$ of PTF10ngv’s optical position, suggesting that it is unlikely to be a counterpart to the star.

3.2. Spectroscopic Analysis

PTF10ngv exhibits a rich set of emission-lines, including hydrogen and both low- and high-excitation metallic atomic lines. Atomic line-strength differences between the multiple

optical and IR observations are presented quantitatively in Tables 5 and 6. Equivalent widths (EWs) and line fluxes reported in these tables are calculated by integrating across the line after subtracting off the local continuum, which is estimated by interpolating between two line free regions well separated from the feature in question (typically $v > 1000 \text{ km s}^{-1}$ away). Line fluxes are simply calculated as the direct integral of the line in the flux calibrated spectra; EWs are calculated by dividing the integrated flux by the mean value of the local continuum.

In addition to the many atomic features in emission in PTF10ngv’s spectrum, there is also strong molecular emission from TiO and VO at red optical wavelengths (see PTF10ngv’s discovery spectrum, shown in Figure 7), and from VO, H_2O , and CO in the IR regime (see PTF10ngv’s SpeX and TripleSpec spectra, shown in Figure 8). To the best of our knowledge, this is the first time such prominent molecular emission has been detected from a young star at optical wavelengths.

The four optical spectra from Keck, Lick, and Palomar are compared over the [O I] to Ca II triplet region in Figure 9. Note that the spectra were obtained after the first brightness peak in the 2010 light curve, near the beginning and the end of the $m_R \approx 16.5 \text{ mag}$ nadir (Keck and Lick) and then when the object was brightening again and had reached $m_R \approx 14.8 \text{ mag}$ for the second time during our monitoring. There is some evidence for a change in the overall continuum slope during this month; as the object brightened by several magnitudes it also became redder in a manner similar to interstellar extinction by $\sim 1 \text{ mag}$ in the blue and $\sim 1.5 \text{ mag}$ in the red. The emission feature strengths also varied, as discussed below. The broad TiO band emission, also discussed below, appeared strongest in the July 19 (Lick) spectrum just before the object began its second brightening.

We analyze spectroscopic features of particular interest in more detail below. To analyze the kinematic structure of the observed line profiles, we assume a rest velocity for PTF10ngv of $V_{\text{LSR}} \sim 4 \text{ km s}^{-1}$, consistent with ^{13}CO observations of the surrounding molecular gas (Dobashi et al. 1994).

3.2.1. H I Lines

PTF10ngv has strong H I emission lines including $\text{H}\beta$, $\text{H}\alpha$, the Paschen series (detectable up to 21), and the Brackett series

Table 5
Optical Line Strengths

Line	Wavelength (Å)	LRIS EqW (Å)	LRIS Flux (10^{-15} erg cm $^{-2}$ s $^{-1}$)	Kast (July) EqW (Å)	Kast (July) Flux (10^{-15} erg cm $^{-2}$ s $^{-1}$)	DoubleSpec EqW (Å)	DoubleSpec Flux (10^{-15} erg cm $^{-2}$ s $^{-1}$)
H η	3835.38	3.68	0.087	4.15	0.21
H ζ	3889.05	2.57	0.14	6.57	0.30
Ca II K (abs)	3933.66	1.41	0.056	4.01	0.25
Ca II K (em)	3933.66	-5.65	0.24	-10.57	0.449
H ϵ + Ca II H	3970.07	4.93	0.21	7.04	0.366
H δ	4101.74	2.17	0.12	5.34	0.306
H γ	4340.47	4.00	0.33	7.13	0.57
H β (abs)	4861.33	2.88	0.46	3.92	0.55
H β (em)	4861.33	-1.7	0.27	-0.3	0.04
Na D	5889.95+5895.92	9.637	4.35	7.82	2.85
Fe I	6191.56	-0.61	0.368	-1.82	0.89
[O I]	6300.30	-9.14	5.48	-7.91	3.99
[O I]	6363.78	-3.8	2.245	-2.95	1.41	-1.5	2.22
Fe II	6432.68	-2.11	1.33	-3.38	1.59	-1.68	2.73
Fe I	6495.74	-1.158	0.80	-2.27	1.17	-1.78	2.97
Fe II	6516.08	-2.13	1.5	-2.84	1.49	-2.92	4.89
H α	6562.85	-39.9	31.5	-56.5	31.33	-24.7	48.6
[N II]	6592	-0.86	0.70	-0.18	0.13	-1.15	2.23
[S II]	6716.44	-0.34	0.29	1.26	0.89
[S II]	6730.82	-0.80	0.67	0.82	0.58
Ca I	7148.15	-1.83	2.04	-1.76	1.64
[Fe II]	7172.00	-0.59	0.66	0.58	0.56
[Ca II]	7291.47	-1.08	1.23	-1.63	1.46	-1.14	3.47
[Ca II]	7323.89	-1.02	1.13	-1.12	0.97	-1.19	3.6
K I	7664.91	3.02	4.22	2.24	2.66	3.16	11.1
K I	7698.96	2.32	3.32	2.86	3.80	1.54	5.74
O I	7773	2.26	3.42	2.43	3.35	1.77	7.32
O I	8446.36	-1.52	2.86	-2.47	3.98	-2.16	9.81
Ca II	8498.02	-16.8	32.8	-23.46	35.57	-19.0	85.6
Ca II	8542.09	-19.1	35.9	-24.24	37.63	-16.2	69.3
[Fe II]	8616.952	-1.14	2.1	-0.73	1.11
Ca II	8662.14	-15.4	29.0	-24.88	36.7	-17.55	60.8
H I (Pa 11)	8863.4	-1.47	2.64	-3.14	4.14
H I (Pa 10)	9015.6	-1.58	2.98	-2.15	2.80
H I (Pa 9)	9229.7	-1.61	3.46	-2.73	4.26
H I (Pa 8)	9545.97	-1.44	3.36	-3.59	5.19
[Fe I]	9998.33	-2.28	5.21
H I (Pa 7)	10049.4	-1.82	4.89	-6.35	9.3

including Br- α at 4.05 μ m. No Balmer jump is apparent, but the pileup of the higher-order Brackett lines is visible in the H band of the SpeX observations. The high-dispersion NIRSPEC spectrum reveals profile widths in the Pa- γ and Pa- δ H I lines of ~ 125 km s $^{-1}$ FWHM. The lines are slightly asymmetric with respect to their peak emission at roughly zero velocity, having somewhat more integrated flux on the blue side of line center than on the red side. The H α line EWs are -26 to -46 Å among the four optical spectra, moderate but not extreme relative to Taurus Class I or Class II objects (White & Hillenbrand 2004).

We have derived an estimate of the extinction towards PTF10ngv from the strength of its Pa δ and Br γ emission lines. These lines, representing the 7-3 and 7-4 transitions respectively, share the same upper level. Neglecting collisional effects (a questionable assumption that we will return to below), the relative frequency of radiative de-excitations from the same upper level should be determined entirely by atomic physics. That is, the intrinsic flux ratio of these lines can be calculated in the collisionless limit as

$$\frac{F_{Br\gamma}}{F_{Pa\delta}} = \frac{A_{Br\gamma}\lambda_{Pa\delta}}{A_{Pa\delta}\lambda_{Br\gamma}}, \quad (1)$$

where F , A , and λ give the flux, Einstein coefficient, and wavelength of each line. Using line constants from the NIST atomic database, we calculate an intrinsic Br γ -to-Pa δ line ratio of 0.42 for a collisionless emission region. Densities in magnetospheric accretion columns, however, are sufficiently high that collisions are likely to be an important factor for the emergent fluxes of H I emission lines. Detailed radiative transfer models of relatively cool ($T_{\max} = 6000$ K) magnetospheric accretion columns nonetheless return a Br γ -to-Pa δ line ratio very close to this naive calculation. Models of significantly warmer ($T_{\max} = 9000$ K) accretion columns return a larger line ratio of 2.21 (R. Kurosawa 2010, private communication). The line strength ratio in the SpeX spectrum of PTF10ngv is ~ 8.2 , demonstrating the presence of significant extinction relative to either accretion model. We therefore estimate the extinction to the H I line emission region (which may or may not be coincident with the stellar photosphere) by assuming a standard $R_V = 3.1$ extinction law, and de-reddening the observed line fluxes to match each of these predicted intrinsic line ratios. These estimates place crude bounds on the extinction to PTF10ngv's emission line region of $6 < A_V < 12.4$.

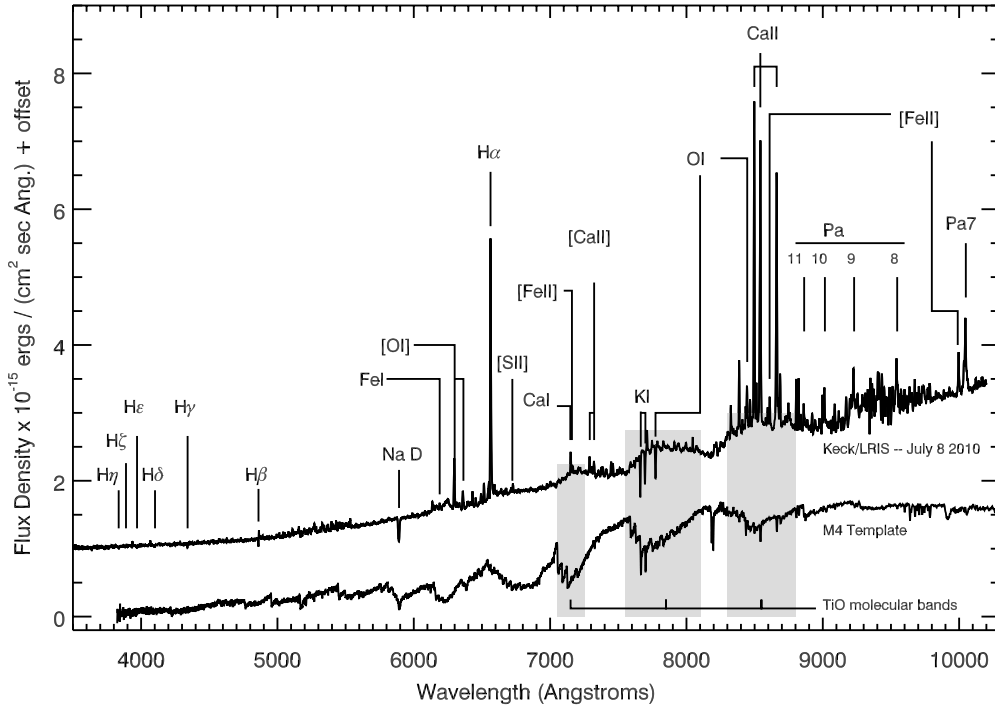


Figure 7. Optical spectrum of PTF10nv obtained with Keck/LRIS on 2010 July 8. Strong emission lines are identified, with line strengths presented quantitatively in Table 5. An M4 field-dwarf spectral template (Bochanski et al. 2007) is shown to illustrate the correspondence between the TiO emission bands in PTF10nv and typical TiO absorption bands in cool M-type photospheres.

Table 6
Near-infrared Line Strengths

Line	Wavelength (μm)	SpeX EqW (\AA)	SpeX Flux ($10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$)	TripleSpec EqW (\AA)	TripleSpec Flux ($10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$)
Ca II	0.8662	-12.91	-26.4
Ca II	0.8542	-13.49	-28.5
Ca II	0.8499	-13.97	-30.73
H I 7-3	1.0049	-0.698	-1.706	-4.04	-7.10
H I 6-3	1.0938	-3.14	-8.625	-16.01	-19.4
O I	1.1287	-4.211	-12.397	-13.01	10.06
Fe I	1.1595	-0.733	-2.293	-1.15	-2.29
Fe I	1.161	-1.294	-4.152	-3.30	-6.83
K I	1.169	-1.449	-4.539	-3.13832	-6.80
Fe I	1.179	-0.807	-2.665	-1.45	-3.55
Mg I	1.1833	-2.092	-6.881	-2.64	-6.16
Fe I	1.1885	-3.018	-10.106	-5.26	-12.47
Fe I	1.1974	-1.847	-6.444	-3.48	-8.89
H I 5-3	1.2818	-7.766	-31.026	-16.5	-50.4
Mg I	1.504	-3.777	-23.54	-5.39	-29.7
Si I	1.589	-2.372	-16.094	-2.47	-15.3
Mg I	1.711	-0.587	-4.73	-0.73	-5.50
H ₂ (1-0 S(1))	2.121	-0.848	-10.12	-1.35	-16.9
H I 7-4 (Br γ)	2.1657	-1.186	-14.40	-2.24	-28.58
Na I	2.2075	-0.768	-9.53	-0.957	-12.49
H ₂ (2-1 S(1))	2.247	-0.199	-2.412	-0.41	-5.52
Ca I	2.282	-0.503	-6.5	-1.07	-14.9
CO	2.2923	-18.8	-249.9	-29.3	-423.9

Methods have also been developed to estimate a young star's mass accretion rate from the strength of its NIR H I emission lines (e.g., Muzerolle et al. 1998). These relationships, which link observed H I line luminosities with mass accretion rates inferred from UV continuum excesses, have only been calibrated for observations of T Tauri stars; applying these relationships to more heavily embedded protostars requires the assumption that the dynamics of the underlying accretion process are similar

across these evolutionary phases, and that the line luminosity can be measured accurately even when the protostar's photosphere is not detectable. These assumptions are even more problematic for an extreme accretor such as PTF10nv, but an accretion rate estimate via H I line fluxes can still provide a useful lower limit on the object's accretion rate: changes in the accretion dynamics (i.e., boundary-layer accretion instead of magnetospheric accretion) or the presence of significant veiling

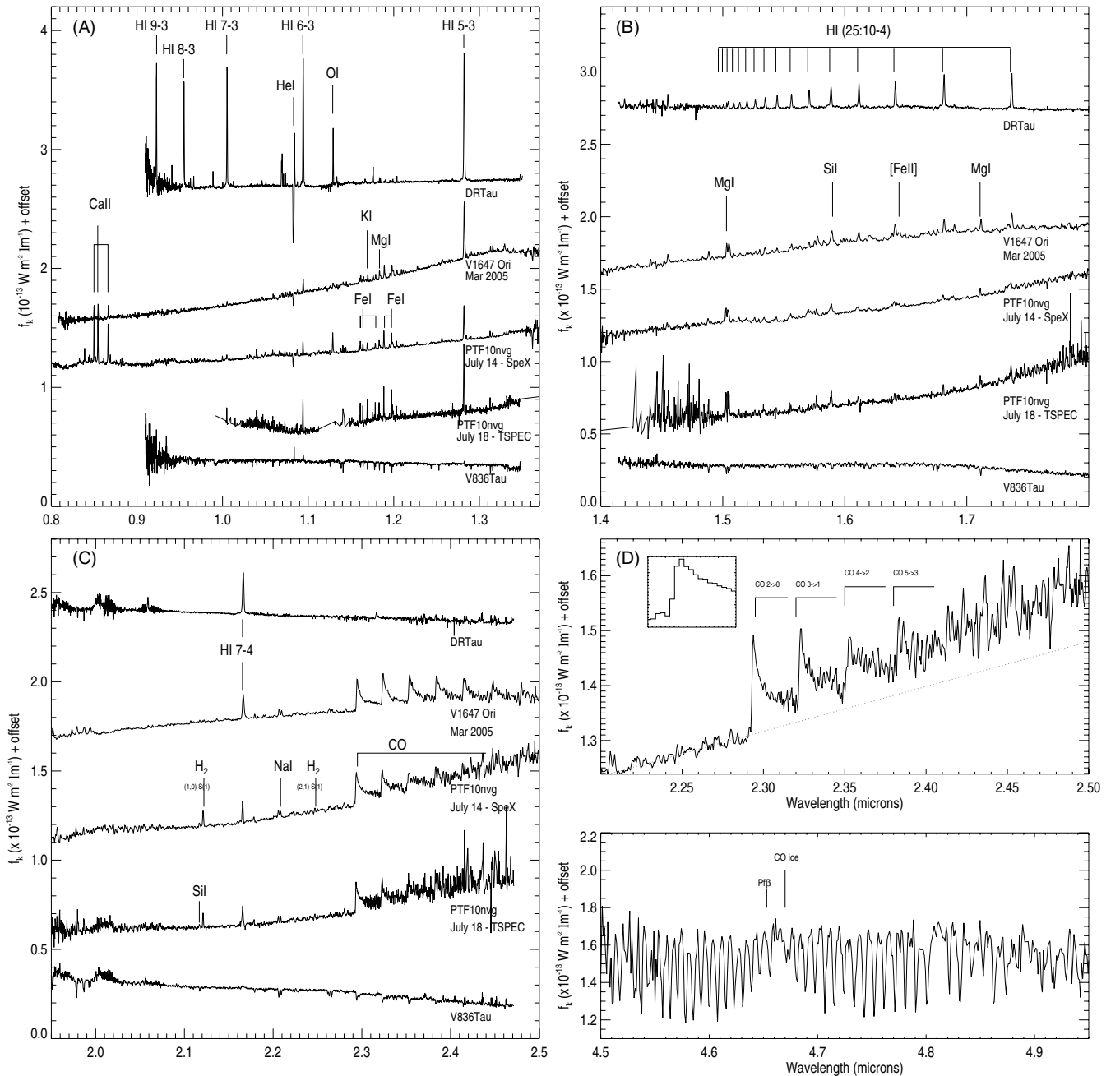


Figure 8. *YJ*-, *H*-, *K*-, and *M*-band spectra of PTF10nvg in panels (A), (B), (C), and (D), respectively. Panels (A)–(C) also include spectra of DR Tau (heavily accreting Class II object), V836 Tau (typical Class III object), and V1647 Ori (McNeil’s Nebula outburst) for comparison. Strong emission lines are labeled where possible; PTF10nvg line strengths are also presented quantitatively in Table 6. Notably, the EWs of the emission lines in PTF10nvg are 50%–100% larger in the APO TripleSpec observation than in the SpeX observation even though the line fluxes are similar. The difference in EWs therefore likely reflects differences in the continuum flux level between the two epochs.

flux will both likely result in an underestimate the object’s full line flux, and/or an increase in the conversion factor between line luminosities and accretion rates.

We have generated crude mass accretion rate estimates for PTF10nvg from the strengths of its Br- γ and Pa- β emission lines. We have dereddened the Pa- β and Br- γ line fluxes listed in Table 6 assuming an $A_V \approx 9$ mag (intermediate to the bounds calculated above), and converted those dereddened fluxes into line luminosities assuming a distance to PTF10nvg of ~ 600 pc. Using Equations (1) and (2) from Muzerolle et al. (1998) to estimate $\text{Log } L_{\text{acc}}/L_{\odot}$ from each line, we infer $\text{Log } L_{\text{acc}}/L_{\odot} = 0.62$ and 0.335 from the Pa- β and Br- γ lines, respectively.

Assuming canonical Classical T Tauri Star parameters ($M = 0.5 M_{\odot}$ and $R = 3 R_{\odot}$), these accretion luminosities imply mass accretion rates of $\dot{M} \approx 2.5 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$.

3.2.2. Ca II

The Ca II triplet lines at 8498, 8542, and 8662 Å are also generally associated with accretion activity in young stars. PTF10nvg exhibits EWs for these lines of -15 to -25 Å, varying among the observations, but in the realm of the rapidly accreting young Class II stars in Taurus and comparable to the emission levels measured by Connelley & Greene (2010) in a large survey of Class I objects. The line ratios seen

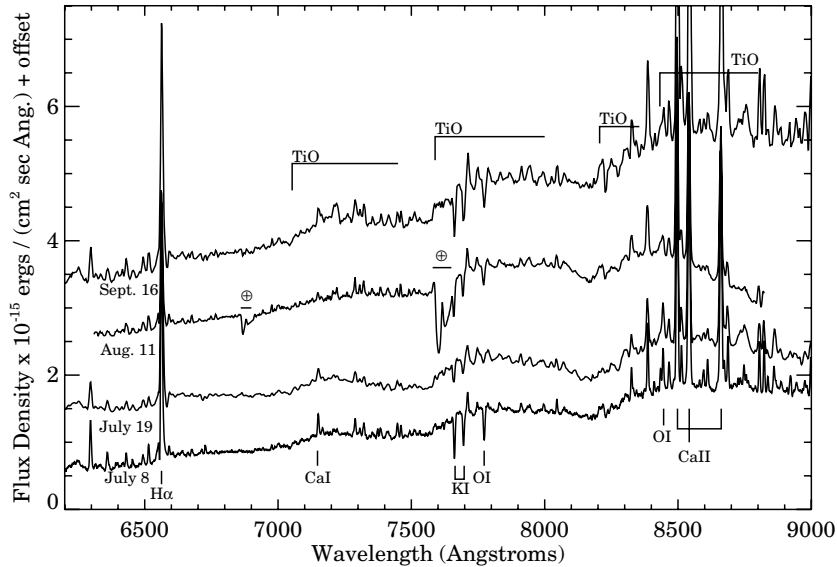


Figure 9. Comparison of the four optical spectra between the [O I] and the Ca II triplet region (note that the August 11 Palomar DBSP spectrum has not been corrected for telluric absorption). The apparent “continuum” differences are due to variations in the strength of the TiO emission over the four epochs.

from PTF10ngv varied, being somewhat typical of Class II sources ($F_{\text{Ca II } 8498}/F_{\text{Ca II } 8542} = 0.91$, $F_{\text{Ca II } 8662}/F_{\text{Ca II } 8542} = 0.81$) in the LRIS data taken on 2010 July 8, but roughly equal ($F_{\text{Ca II } 8498}/F_{\text{Ca II } 8542} = 1.08$, $F_{\text{Ca II } 8662}/F_{\text{Ca II } 8542} = 0.93$) in the SpeX data from July 14. The latter case is again similar to heavily accreting T Tauri stars at the extreme of the Class II range illustrated by Hamann & Persson (1992) and close to the ratio predicted by their highly optically thick ($\tau = 1000$) slab models. It is also similar the outbursting behavior of V1647 Ori (Walter et al. 2004).

Like the Paschen lines, the Ca II triplet lines appear slightly asymmetric with respect to their line centers, having somewhat more flux on the blueshifted sides of the lines than the red.

The Ca II K line is also in emission, though the H line is affected by blueshifted Balmer line absorption.

3.2.3. Other Atomic Lines

The only absorption features we detect in these optical spectra are blueshifted lines of H I, Na I D, and K I, along with the O I triplet at 7773 Å, which appears centered at the systemic velocity. The IR spectra, similarly, have only a single blueshifted line in absorption (the He I λ 10830 feature seen in Figure 8 and discussed in Section 3.2.4), so no spectral type can be derived for the outbursting object. Instead of a stellar photosphere, the spectrum is dominated by the circumstellar signatures of accretion and outflow processes.

In addition to the H I and Ca II lines discussed in Sections 3.2.1 and 3.2.2, the optical and NIR spectra of PTF10ngv show many emission lines characteristic of young accreting stars. The optical spectra show several permitted and forbidden species, including O I, Fe I, Fe II, Ca I, Ca II, [Fe II], [Ca II], and [S II]. The high-dispersion Keck NIRSPEC Y-band spectrum shows low-excitation (<1 – 2 eV) Ti I lines as well as higher excitation (6–8 eV) Si I and C I lines in abundance, identified from the line list presented by Sharon et al. (2010). A Y-band survey including many Class II and some Class I stars in Taurus, as well as other young stars (Edwards et al. 2006; Fischer et al. 2008), suggests that some extreme emission line objects display the higher excitation lines. The Ti I is quite rare, however, and is found in our experience only in the spectra of V1331

Cyg, ZCMA, and SVS 13—all embedded sources driving strong outflows. The lower dispersion SpeX spectrum shows emission lines from many neutral species including Al I, Ca I, Fe I, K I, Mg I, Na I, O I, Si I, and Ti I over the 1–2.5 μm region; many of these features are also visible in the NIR outburst spectrum of V1647 (shown for comparison in Figure 8).

3.2.4. Wind Indicators: Blueshifted H I, Na D, K I, and He I λ 10830 Absorption and [O I] Emission

There is significant evidence for a substantial wind associated with the outburst of PTF10ngv. Our optical spectra have limited resolution of ~ 150 – 300 km s^{-1} , but nevertheless the H α and H β lines display clear P Cygni features with sub-continuum absorption out to -400 km s^{-1} (see kinematic profiles shown in Figure 10). The upper Balmer series lines lack the redshifted emission and display only a broad, blueshifted absorption trough. In the NIR, where our NIRSPEC spectrum provides a higher resolution by a factor of 2–3, none of the Paschen or Brackett lines has P Cygni structure. The Na I D lines, like the upper Balmer lines, are seen in blueshifted absorption out to about -400 km s^{-1} but are blended at our resolution.

The K I doublet at 7664 and 7698 Å is seen in absorption in PTF10ngv’s spectrum, blueshifted by ~ 175 km s^{-1} ; this line is not typically seen in the spectra of Class I or Class II stars. Strong and blueshifted K I and Na I D absorption, as well as P Cygni Balmer profiles, are seen in the optical spectra of the FU Ori stars V1057 Cyg and V1515 Cyg, however, and in the outburst spectrum of V1647 Ori.

The high-dispersion NIRSPEC spectrum shows a somewhat complex He I λ 10830 profile (see Figure 11). Not unlike many heavily accreting T Tauri stars, such as DR Tau (Edwards et al. 2003, 2006), there is strong blueshifted absorption out to about -300 km s^{-1} which reaches 60% of the continuum level. The absorption does not extend continuously inward to zero velocity, however; instead, the continuum level is reached by about -100 km s^{-1} and there is a second narrow (FWHM ≈ 17 km s^{-1}) absorption component centered at PTF10ngv’s rest velocity. Similar low-velocity and narrow absorption components are seen in the weakly or nonaccreting T Tauri stars V836

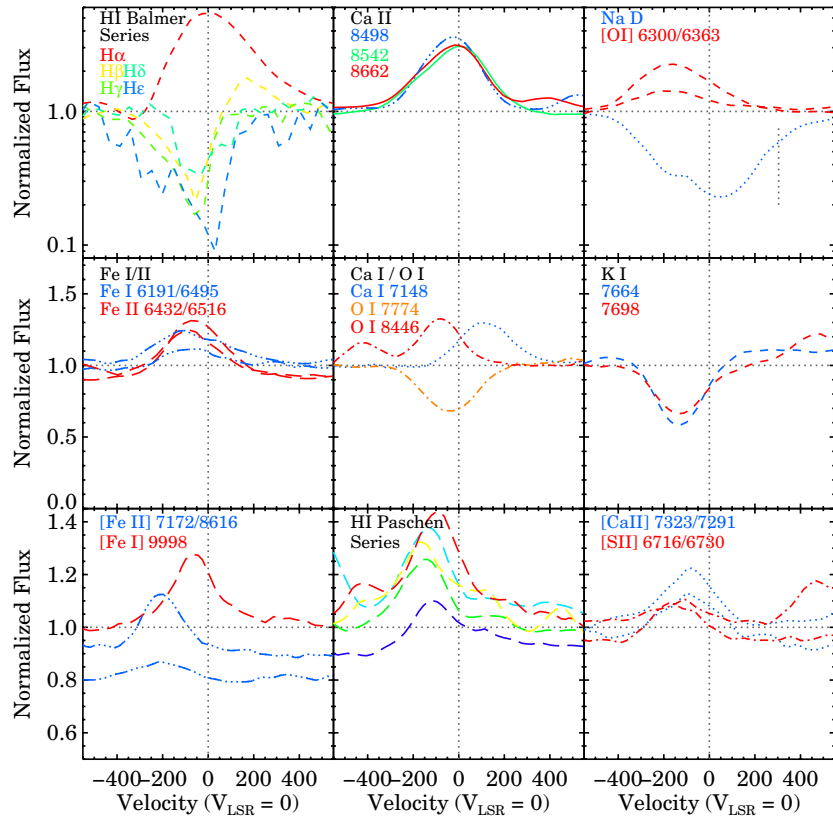


Figure 10. Selected spectral line profiles from the Keck LRIS spectrum, displayed with a velocity scale relative to the rest velocity of the Pelican Nebula ($V_{\text{LSR}} \approx 0 \text{ km s}^{-1}$).

(A color version of this figure is available in the online journal.)

Tau and TWA 2. When considering both the broad absorption terminal velocity and the narrow absorption depth, the He I $\lambda 10830$ profile of PTF10ngv is quite similar to that of V1057 Cyg (unpublished data). The origin of a narrow absorption feature at roughly rest relative to the star is unknown; we note, however, the similar feature of V836 Tau was attributed to a disk wind by Kwan et al. (2007). Also of note is the possible presence of Si I emission at 10830 \AA that could be filling in some of the blueshifted He I absorption at intermediate velocities; this can be gauged by comparison to the Si I $\lambda 10869$ line also shown in Figure 11.

Other typical indicators of mass loss (Hartigan et al. 1995) are the [O I] $\lambda\lambda 6300, 6364$ emission lines, which in PTF10ngv are symmetric and blueshifted by $\sim 175 \text{ km s}^{-1}$, and the [S II] $\lambda\lambda 6717, 6731$ lines, also blueshifted by $\sim 175 \text{ km s}^{-1}$. The [N II] $\lambda\lambda 6548, 6584$ lines are relatively weak and [O I] $\lambda 5577$ is not apparent in our low spectral resolution data. The forbidden-line EWs are comparable to those seen in the some of the strongest Class II stars in Taurus such as CW Tau, DD Tau, DG Tau, HN Tau, and UZ TauE (Hartigan et al. 1995), but weaker than those typical of Class I and other stars in Taurus driving strong outflows (White & Hillenbrand 2004).

Hartigan et al. (1995) used detailed shock models to derive relations allowing a young star's mass loss rate to be estimated from the luminosities of its [O I] 6300 and [S II] 6731 forbidden lines. These relations are based upon assumptions about the typical density, velocity, and abundance of the shock region, and also require that observed line fluxes be converted to line luminosities, with corrections to account for reddening. Given the strong outflow signatures PTF10ngv possesses, it is possible, and perhaps likely, that the parameters commonly

assumed for T Tauri star jets are not a good description of the physical properties of shocked regions in PTF10ngv's outflow. Nonetheless, we can convert the observed [O I] 6300 and [S II] 6716 line fluxes to luminosities by adopting the same distance and extinction estimates used above to calculate H I line luminosities, and thus infer crude estimates of PTF10ngv's mass loss rate from Hartigan, Edwards & Ghandour's (1995) equations A8 and A10. These calculations imply estimated mass loss rates of $7 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$ (as inferred from [S II] 6716) and $2 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ (as inferred from [O I] 6300). Given the sizable uncertainties in the adopted extinction correction, and even the appropriateness of the physical properties assumed for the shock regions, these mass loss rates may easily be incorrect by more than an order of magnitude. Nonetheless, they do emphasize the considerable mass loss associated with the PTF10ngv outburst; relative to the $\sim 10^{-7} M_{\odot} \text{ yr}^{-1}$ accretion rate inferred above from the H I emission lines, these estimates suggest an inflow/outflow ratio $\gtrsim 1$, one to two orders of magnitude larger than the 0.01–0.1 ratios typically inferred for T Tauri stars.

NIR line ratios provide additional evidence that PTF10ngv is driving a strong outflow. Lorenzetti et al. (2009) and Antonucci et al. (2008) have noted that EXor/jet driving sources typically demonstrate flux ratios between their Br- γ and $2.206 \mu\text{m}$ Na I emission lines of ~ 2 –5. Class I sources that lack jet signatures, however, tend to have ratios that favor Br- γ much more strongly (e.g., 6–20; Nisini et al. 2005). The NIR spectra of PTF10ngv presented here possess a ratio of these lines of ~ 2 , placing PTF10ngv firmly in the company of the jet driving sources, as is consistent with the morphology of the wind-sensitive lines analyzed above.

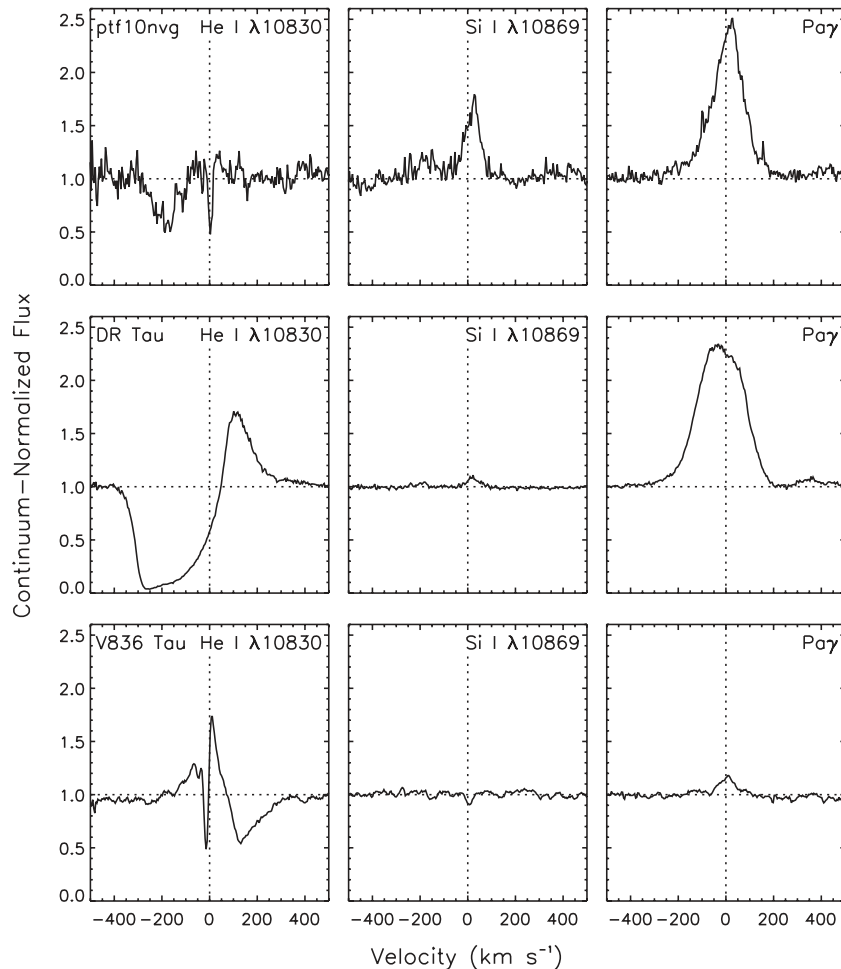


Figure 11. Keck NIRSPEC profiles of He I $\lambda 10830$, Si I $\lambda 10869$ in order to assess the possible contamination of Si I $\lambda 10830$ to the He I profile, and the H I Pa- γ line of PTF10nvg. DR Tau and V836 Tau are shown for comparison: the two Taurus stars are benchmarks for their comparable terminal velocity in the broad blueshifted absorption (DR Tau) and the roughly systemic narrow absorption (V836 Tau) components to the He I profile.

3.2.5. Outflow as Diagnosed Via H_2 and $[Fe II]$ Emission

PTF10nvg shows clear emission in the $2.12 \mu\text{m } H_2(1-0)S(1)$ line, as well as a weaker feature which is likely the $2.24 \mu\text{m } S(2-1)S(1)$ line. Modeling has established that the ratio of these line strengths can diagnose between collisional and radiative excitation mechanisms: pioneering work by Gredel & Dalgarno (1995) demonstrated that UV and X-ray irradiation will produce distinctive $H_2(1-0)S(1)/S(2-1)S(1)$ line ratios (0.54 and 0.06, respectively), and could also be distinguished from the ratio expected from a $T = 2000 \text{ K}$ shocked gas (0.13). More recently, Nomura et al. (2007) calculated synthetic H_2 spectra as expected to emerge from protoplanetary disks with realistic gas-to-dust ratios and density/temperature profiles. The Nomura et al. (2007) models predict that H_2 excitation by UV radiation is sensitive to the dust properties in the disk, producing $H_2(1-0)S(1)/S(2-1)S(1)$ line ratios from 0.025–0.23 for power-law distributions of dust grains with maximum sizes of $10 \mu\text{m}$ to 10 cm , respectively. X-rays are much less sensitive to dust properties; models where X-rays dominate the excitation produce ratios of ~ 0.07 , irrespective of the dust properties.

The $H_2(1-0)S(1)$ and $S(2-1)S(1)$ line strengths recorded in Table 6 correspond to ratios of 0.23 and 0.3 for the SpeX and TSPEC observations; the less secure measurement of the $(1-0)S(1)$ line suggest this ratio may best be treated as a tenuous detection which could be uncertain at the factor of two level.

Even accounting for this relatively large uncertainty, models attributing the excitation to X-rays alone are essentially ruled out; models incorporating UV or UV + X-ray heating and dust grains $\lesssim 1 \text{ mm}$ would also be difficult to reconcile with this ratio. The observed values are most consistent with, though slightly larger than, the Gredel & Dalgarno (1995) model of shock excitation or the Nomura et al. (2007) model of UV + X-ray irradiation of a disk with dust subject to coagulation and settling, which predict $H_2(1-0)S(1)$ and $S(2-1)S(1)$ line ratios of 0.13 and ~ 0.18 , respectively. Given the wealth of outflow signatures present in PTF10nvg’s spectrum, we consider shocks to be the most likely source, but this conclusion is far from established.

$[Fe II]$ is another common outflow tracer for young stars: Connelley & Greene (2010) demonstrate that the line strengths of H_2 and $[Fe II]$ are strongly correlated for a large sample of Class I protostars. The TripleSpec observation of PTF10nvg shows clear $[Fe II] \lambda 1.257$ and $\lambda 1.644 \mu\text{m}$ emission; these lines may also be present in the SpeX observation, but the lower spectral resolution leads to a lower line-continuum contrast. In comparison to the Class I stars in the Connelley & Greene (2010) sample, PTF10nvg shows relatively weak H_2 and $[Fe II]$ emission. Connelley & Greene (2010) note, however, that the stars in their sample with weak H_2 also have low levels of continuum veiling, which is not true for PTF10nvg, where we cannot detect any photospheric absorption features.

As both the 1.257 and 1.644 μm lines share the same upper level, Connelley & Greene (2010) use their observed line fluxes to infer extinction estimates for each star in their sample, as with the H I-based estimates presented in Section 3.2.1. We find a line ratio in PTF10ngv of 1.48, essentially equal to that assumed by Connelley & Greene (2010) for an unextinguished source, in strong disagreement with the extinction estimate of 6–12 which we infer from the H I lines of PTF10ngv. Connelley & Greene (2010) find a similar discrepancy, however, in their analysis of the Class I source 06297+1021W, whose H I lines imply $A_V \approx 18$ mag but whose [Fe II] lines suggest $A_V \approx -1$ mag. In many regards, 06297+1021W is quite similar to PTF10ngv: in addition to their similarly incongruous H I and [Fe II] extinction estimates, both show weak H₂ and [Fe II] emission relative to other Class I sources (Connelley & Greene 2010 are unable to measure values for these features for 06297+1021W), both objects appear as a *K*-band point source absent significant nebulosity, and both display a rich, morphologically similar NIR emission-line spectrum.

Connelley & Greene (2010) also explore the correlation between the strengths of the H₂, [Fe II], and Br- γ emission lines: relative to their sample of Class I objects, PTF10ngv lies on the weak side of these relations, though not significantly separated from the rest of the Class I population.

3.2.6. Optical TiO/VO Emission

In addition to the rich atomic emission-line spectrum, the most salient feature of the optical spectrum of PTF10ngv is the broad emission that we attribute to TiO and VO. The PTF10ngv spectrum (Figure 7) looks much like an inverted mid-M spectral type star, suggesting that in addition to the hot and moderate density emission lines, there is a cooler (~ 1500 – 4000 K, Lodders 2002; Ferguson et al. 2005; Sharp & Burrows 2007) and high-density ($> 10^{10}$ cm⁻³) emitting component, perhaps the outer layers of a disk. Notably, the TiO emission is variable among our three spectra, being strongest just before the object brightened for the second time during the 2010 season.

To our knowledge, this is the first time such prominent hot molecular emission has been reported for a young star. There is, however, literature discussion of emission in the optical TiO band heads, though not the broader full bands as we observe in PTF10ngv. TiO band head emission has been mentioned in the spectra of several luminous B[e] stars (Zickgraf et al. 1989) and several red giant or supergiant dusty objects, notably VY CMa (Phillips & Davis 1987) and U Equ (Barnbaum et al. 1996), as well as in the red nova V4332 Sgr (Goranskii & Barsukova 2007). Most relevant to the present case is the report by Herbig (2009) that V1057 Cyg exhibited emission in 2004 in “the heads of the (0-0) bands of the γ system of TiO at 7054, 7087, and 7125 Å.” We have confirmed this narrow emission in our own data on V1057 Cyg but emphasize that the PTF10ngv TiO emission is through the full band region, not just in the band heads.

3.2.7. Infrared CO, VO, and H₂O Emission Lines

In addition to the optical TiO and VO emission, there is also molecular emission throughout the 1–5 μm spectra of PTF10ngv. Readily apparent broad features above the likely dust-dominated red continuum are VO at 1.05 μm , several H₂O bands around 1.4, 1.8, 2.0, and 2.45 μm as well as the Q and R branches at 3.0 and 3.3 μm , and prominent CO first overtone band heads beginning at 2.3 μm . While the VO and H₂O suggest emitting region temperatures of < 3500 K, the CO emission

could come from warmer (2500–5000 K) dense ($> 10^{10}$ cm⁻³) gas.

Connelley & Greene (2010) display correlations between the morphology and strengths of the CO, Na, Ca, and Br- γ features for Class I sources, finding that all sources with CO in emission also display Br- γ and Na in emission. PTF10ngv follows this relation and displays CO, Na, and Ca emission-line strengths similar to those of other Class I sources with each set of features in emission. PTF10ngv does show somewhat weaker Br- γ relative to its CO emission strength than most Class I sources, but is not an extreme outlier to that relation.

In contrast to the CO band head emission at 2.3 μm , the 4.7 μm CO fundamental band is in *absorption* with both the P and R branches evident.

4. DISCUSSION

The spectra obtained to date of PTF10ngv lack the strong FG-supergiant optical and M-supergiant NIR absorption features typical of FU Ori objects, which are interpreted as disk-atmosphere dominated systems. Instead, the PTF10ngv spectra are characterized mainly by emission lines, some with kinematic signatures of formation in outflowing material. The only similarities to the FU Ori class are in the Na I D, K I, and He I blueshifted absorptions which are kinematically associated with strong outflows, as well as the O I 7773 Å neutral velocity absorption which also may be related to winds. Despite the optical rise by over 6 mag within one year, PTF10ngv does not appear to be in the FU Ori class. This statement is based on the first year of data on this object and does not preclude that it could settle down in terms of its variability and its outflow-dominated optical and NIR spectrum to reveal FU Ori-like absorption features in the future. There is, however, a notable parallel to the FU Ori object V1057 Cyg in the He I $\lambda 10830$ profile and a similarity in the optical TiO emission.

The similarities of PTF10ngv to EX Lup outbursts are also slim. Neither the light curve nor the spectra are similar to this class of objects, which are interpreted as episodes of enhanced accretion but less extreme than those nominally driven by disk instability in FU Ori objects.

Instead, the PTF10ngv data bear qualitative and quantitative similarities to the 2004 outburst of V1647 Ori (also known as McNeil’s Nebula), albeit with a significantly weaker wind component. Specifically, the amplitude of the optical outburst is similar, although PTF10ngv is much more unstable and photometrically variable in its outburst phase (see the *I*-band light curve given by Acosta-Pulido et al. (2007)).

The PTF10ngv SpeX spectrum is a reasonable match to the one-year post-outburst spectrum of V1647 Ori in terms of the emission-line features. The most notable disagreement is the absence of broad 3.3 and 4.7 μm ice features in the spectrum of PTF10ngv: those features are present at all epochs for V1647 Ori and thought to be foreground molecular cloud material not related to the outburst. The early V1647 Ori spectrum (2004 March) also showed P Cygni profiles in Paschen lines, not seen in PTF10ngv, but those profiles disappeared by 2005 March (e.g., Gibb et al. 2006). Along the same lines, compared to only 1.083 μm He I visible in PTF10ngv, V1647 Ori showed strong He I absorption at both 1.083 and 2.05 μm during its recent outburst, with slight red-side emission in both lines (Figure 1 of Vacca et al. 2004), but by 2004 November the 1.083 μm line was much weaker, without red-side emission (Figure 1 of Gibb et al. 2006). Notably, despite the strong signatures of outflow there is

no evidence of shocked H₂ emission in V1647 Ori. PTF10ngv, in contrast, has a relatively slow wind but does possess H₂ in emission. This could be a line-of-sight effect but may be an intrinsic astrophysical difference. There are several objects in the Connelley & Greene (2010) catalog of Class I sources which show somewhat similar NIR spectra to PTF10ngv and also display H₂ in emission: 16289–4449 (another eruptive variable with CO in emission), and 20453+6746, whose spectrum is somewhat more in line with a typical Class I source but which does display some of the same emission lines in the *J* band. 06297+1021W, the Class I source in the Connelley & Greene (2010) catalog whose similarities to PTF10ngv were outlined in Section 3.2.5, does not, however, appear to show strong H₂ emission.

The PTF10ngv optical spectra also appear similar to those of V1647 Ori in terms of line EWs and flux ratios, though they indicate weaker P Cygni structure in the Balmer and Ca II lines (e.g., Reipurth & Aspin 2004a; Briceño et al. 2004), which themselves disappeared in V1647 Ori roughly one-year post-outburst (Ojha et al. 2006).

A unique feature, however, of PTF10ngv relative to any known young stellar object is the molecular TiO and VO seen in emission in the optical spectra. This emission has varied in strength relative to the varying continuum over a 1 month timescale. It indicates the presence of dense warm gas, perhaps the upper levels of a disk atmosphere heated by irradiation from enhanced accretion luminosity.

5. SUMMARY

1. We have identified PTF10ngv, a Class I protostar undergoing a strong outburst. A comparison of the source’s pre- and post-outburst SED suggests that the bolometric luminosity has increased by a factor of more than 100.
2. This outburst is presumably driven by a significant increase in PTF10ngv’s mass accretion rate. We analyzed the strength of the NIR H1 lines to derive an extinction estimate of $A_V \sim 6\text{--}12$ mag and an accretion rate of $\dot{M} \approx 2.5 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$.
3. Several wind-sensitive optical and NIR absorption lines exhibit blueshifts of several hundred km s⁻¹, indicating that PTF10ngv is driving a substantial outflow. This interpretation is supported by the presence of several spectroscopic lines thought to trace shocked gas (H₂, [S II], [Fe II]).
4. Several optical TiO bands are seen fully in emission, a characteristic that is, to the best of our knowledge, unique to this astronomical object. This emission reveals the presence of a significant amount of dense ($n > 10^{10} \text{ cm}^{-3}$), warm (1500–4000 K), presumably circumstellar gas.
5. During its present outburst, PTF10ngv’s photometric and spectroscopic properties differ significantly from those commonly associated with EXor or FU Ori outbursts. We do identify several objects whose NIR spectra show morphological similarities to that of PTF10ngv: 06297+1021W, 16289–4449, and 20453+6746, as well as V1647 Ori, at least during its 2004–2006 outburst. Detailed monitoring of PTF10ngv, and each of these other variables, is required to ascertain the exact nature of these enigmatic sources.

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Facilities: PO:1.2m, PO:1.5m, Hale, Keck:I, Keck:II, FLWO:1.2m, ARC, Shane

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