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### Quantifying hiss-driven energetic electron precipitation: A detailed conjunction event analysis

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## RESEARCH LETTER

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## Key Points:

- Measured and calculated hiss  $B_w$  from POES electron measurements agree well
- Electron ratio measured by POES is able to estimate hiss wave intensity
- This technique can be used to provide global hiss wave distribution

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## Quantifying hiss-driven energetic electron precipitation: A detailed conjunction event analysis

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**Abstract** We analyze a conjunction event between the Van Allen Probes and the low-altitude Polar Orbiting Environmental Satellite (POES) to quantify hiss-driven energetic electron precipitation. A physics-based technique based on quasi-linear diffusion theory is used to estimate the ratio of precipitated and trapped electron fluxes ( $R$ ), which could be measured by the two-directional POES particle detectors, using wave and plasma parameters observed by the Van Allen Probes. The remarkable agreement between modeling and observations suggests that this technique is applicable for quantifying hiss-driven electron scattering near the bounce loss cone. More importantly,  $R$  in the 100–300 keV energy channel measured by multiple POES satellites over a broad  $L$  magnetic local time region can potentially provide the spatiotemporal evolution of global hiss wave intensity, which is essential in evaluating radiation belt electron dynamics, but cannot be obtained by in situ equatorial satellites alone.

### 1. Introduction

Resonant electron interactions with plasmaspheric hiss play an important role in electron precipitation over a broad range of energies. Electron precipitation loss driven by plasmaspheric hiss is energy-dependent with longer lifetimes of 1–100 days for higher energy electrons ( $> \sim 2$  MeV) and shorter lifetimes less than a day for lower energy electrons ( $< \sim 200$  keV) [Meredith *et al.*, 2007; Thorne *et al.*, 2013a; Ni *et al.*, 2013]. Hiss-driven relativistic electron precipitation loss is known to create the slot region between the inner and outer radiation belts [Lyons and Thorne, 1973; Abel and Thorne, 1998] and is also responsible for the decay of the “electron storage ring” between the inner and outer radiation belts, as recently observed by the Van Allen Probes [Baker *et al.*, 2013; Thorne *et al.*, 2013a].

Since plasmaspheric hiss plays an important role in the loss process of energetic electrons in the inner magnetosphere, it is critical to understand the global spatiotemporal evolution of hiss wave intensity during geomagnetic storms. However, in situ satellite measurements of hiss wave intensity in the equatorial magnetosphere are limited to certain  $L$  and magnetic local time (MLT) ranges. Recently, by taking advantage of the extensive coverage of multiple NOAA Polar Orbiting Environmental Satellites (POES), Li *et al.* [2013] adopted a physics-based technique to infer the chorus wave intensity from the low-altitude electron observations (30–100 keV). In the present paper, we apply this technique to a conjunction event between the Van Allen Probes near the equatorial plane and the NOAA 17 satellite at the conjugate low altitudes to quantitatively evaluate the effect of hiss-driven energetic electron scattering near the bounce loss cone.

### 2. Methodology

To quantify hiss-driven electron precipitation, we use quasi-linear diffusion theory [Kennel and Petschek, 1966] to treat the particle scattering due to resonant wave-particle interactions and the University of California, Los Angeles (UCLA) full diffusion code [Ni *et al.*, 2008; Shprits and Ni, 2009] to calculate electron pitch angle diffusion coefficients.

### 2.1. POES Electron Observations

NOAA POES are polar orbiting, Sun-synchronous satellites at altitudes of ~800–850 km with an orbital period of ~100 min. The space environment monitor 2 instrument package onboard POES includes the medium energy proton and electron detector, which has two electron solid-state detector telescopes measuring electron fluxes in three energy bands (>30 keV, >100 keV, and >300 keV) [Evans and Greer, 2004; Green, 2013]. The field of view of the 0° telescope is approximately outward along the local zenith, whereas the 90° telescope is mounted perpendicular to the 0° telescope. Both of the telescopes have fields of view that are ±15° wide. The 0° telescope measures precipitated flux inside the bounce loss cone for  $L > 1.4$  [Rodger et al., 2010a], while the 90° telescope measures mixed particles which may be in the bounce or drift loss cone or stably trapped over the invariant latitude range 55°–68° [Rodger et al., 2010b]. Hereafter we refer to electron fluxes measured by the 0° (90°) telescopes as  $J_0$  ( $J_{90}$ ), respectively. We removed the proton contamination using the correction procedure described by Lam et al. [2010] and also excluded the data within the region of the South Atlantic anomaly to eliminate the drift loss cone effect.

### 2.2. Wave and Electron Measurements From Van Allen Probes

Wave measurements are provided by the Van Allen Probes with a perigee of ~1.1  $R_E$  and apogee of 5.8  $R_E$  in the near-equatorial magnetosphere [Mauk et al., 2012]. The Electric and Magnetic Field Instrument Suite and Integrated Science (EMFISIS) includes a magnetometer instrument which provides measurements of DC magnetic fields and the wave instrument which provides a comprehensive set of wave electric and magnetic fields [Kletzing et al., 2013]. The waveform receiver (WFR) of the waves on the EMFISIS instrument measures wave power spectral density from 10 Hz up to 12 kHz every 6 s and the wave high-frequency receiver (HFR) can measure the upper hybrid frequency, from which the plasma density can be calculated [Kletzing et al., 2013].

Electron fluxes with energy ranging from ~30 keV to ~1000 keV, measured at every 11 s of the spin period by the Magnetic Electron Ion Spectrometer (MagEIS) [Blake et al., 2013], are used in this study to evaluate the electron energy spectrum in the equatorial magnetosphere.

### 2.3. Physics-Based Technique to Link Wave Amplitudes and Electron Precipitation

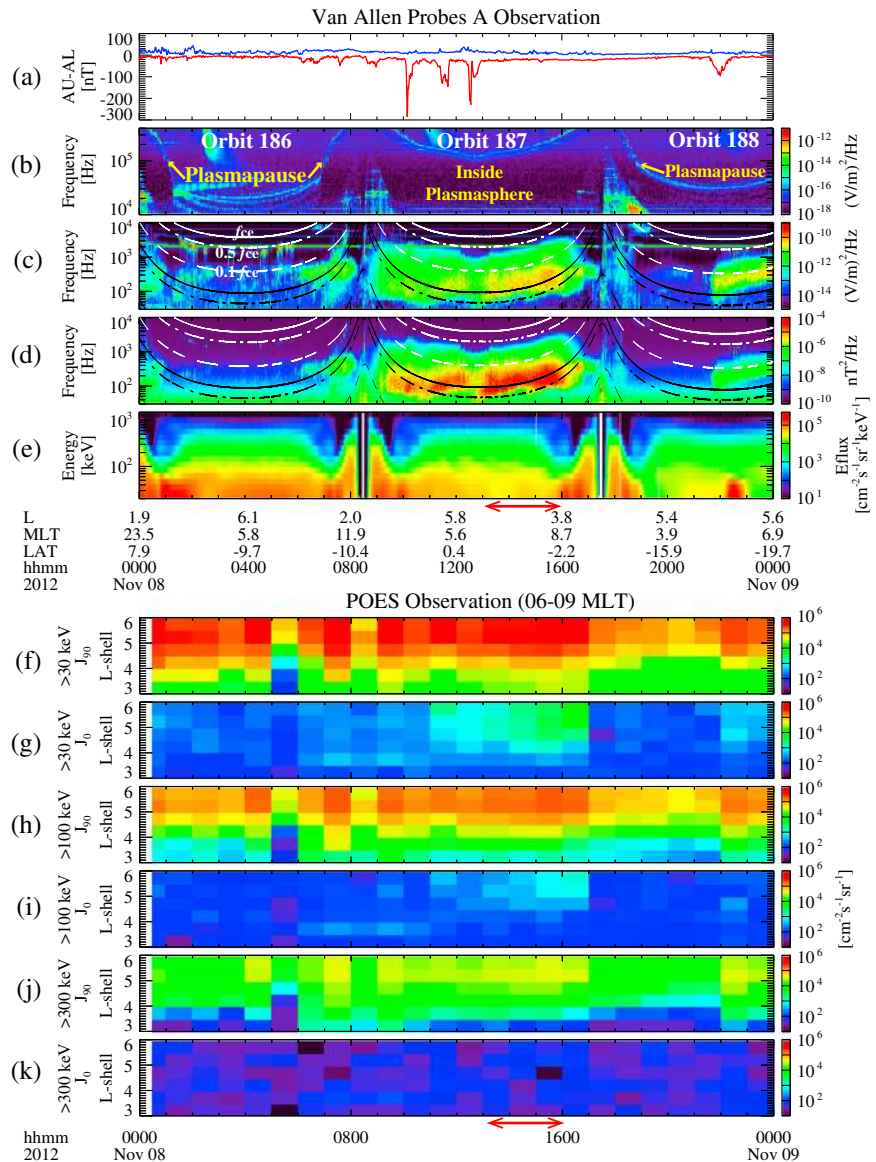
We evaluate the wave-driven electron pitch angle scattering using quasi-linear diffusion theory [Kennel and Petschek, 1966] and link the estimated electron pitch angle distribution to the two-directional POES electron measurements using a newly developed physics-based technique, which has been successfully applied to chorus waves [Li et al., 2013; Thorne et al., 2013b]. The key equation that connects the ratio ( $R$ ) of electron count rates (C.R.) measured by the 0° and 90° telescopes to wave scattering rates is

$$R = \frac{C.R.(E > E_{th})|_{0^\circ}}{C.R.(E > E_{th})|_{90^\circ}} = \frac{\int_{E_1}^{E_2} \int_0^\beta \frac{J(E)}{D_{aa}(E)|_{\alpha=\alpha_0} \cos \alpha_0} \frac{I_0\left(\frac{\alpha_{eq,in}}{\alpha_0} z_0\right)}{z_0 I_1(z_0)} \sin \eta d\eta d\psi dE}{\int_{E_1}^{E_2} \int_0^\beta \frac{J(E)}{D_{aa}(E)|_{\alpha=\alpha_0} \cos \alpha_0} \left[ \frac{I_0(z_0)}{z_0 I_1(z_0)} + \ln \frac{\sin \alpha_{eq,out}}{\sin \alpha_0} \right] \sin \eta d\eta d\psi dE}, \quad (1)$$

where  $\alpha$  and  $E$  are the local electron pitch angle and energy,  $E_1$  and  $E_2$  are the lower and upper electron energy for integration, and  $\beta$  is the half angle of the detector acceptance (15° for the POES particle detector). Here  $\alpha_{eq,in}$  and  $\alpha_{eq,out}$  are the equatorial pitch angles corresponding to the local pitch angle  $\alpha$  for the 0° and 90° telescopes.  $J(E)$  is the electron energy spectrum,  $D_{aa}(E)|_{\alpha=\alpha_0}$  is the bounce-averaged electron pitch angle diffusion coefficient at the equatorial loss cone  $\alpha_0$ , and  $I_0$  and  $I_1$  are the modified Bessel functions.  $z_0 = \frac{\alpha_0}{\sqrt{D_{aa}(E)|_{\alpha=\alpha_0} \cos(\alpha_0) \tau}}$ , where  $\tau$  is a quarter of the electron bounce period and  $\psi$  and  $\eta$  are defined in

Figure 1b in the study of Li et al. [2013]. Details of the derivation of equation (1) are shown in the study of Li et al. [2013], and we reproduce the key equation here only for completeness.

Using equation (1), we can estimate the ratio ( $R$ ) of  $J_0/J_{90}$ , which could be measured by the two-directional POES particle detectors, based on the calculated values of  $D_{aa}(E)|_{\alpha=\alpha_0}$  for given wave spectral properties. In turn, for a given value of  $R$ ,  $D_{aa}(E)|_{\alpha=\alpha_0}$  can be calculated numerically, since  $D_{aa}(E)|_{\alpha=\alpha_0}$  is the only unknown parameter in equation (1).  $D_{aa}(E)|_{\alpha=\alpha_0}$  is dependent on wave amplitude, electron cyclotron frequency, plasma frequency, wave normal distribution, and wave frequency spectrum [Summers, 2005; Glauert and Horne, 2005]. Therefore, for given wave frequency spectrum and plasma parameters, we can also calculate wave amplitudes from  $D_{aa}(E)|_{\alpha=\alpha_0}$ . Note that in order to calculate wave amplitudes, we need to make assumptions on several parameters



**Figure 1.** Observations of plasmaspheric hiss by Van Allen Probe A and energetic electron precipitation by POES satellites near the postdawn sector on 8 November 2012. (a) AU (blue) and AL (red) index. (b) Frequency-time spectrogram of electric field spectral density in the HFR channel. (c) Frequency-time spectrogram of electric field and (d) magnetic field spectral density in the WFR channel. In Figures 1c and 1d, the white lines indicate  $f_{ce}$  (solid),  $0.5 f_{ce}$  (dot-dashed), and  $0.1 f_{ce}$  (dashed), and the black lines indicate  $f_{LHR}$  (solid) and  $0.5 f_{LHR}$  (dot-dashed), where  $f_{ce}$  and  $f_{LHR}$  represent electron cyclotron frequency and lower hybrid resonance frequency, respectively. (e) Energy spectrogram of electron differential flux observed by MagEIS. (f and g)  $J_{90}$  and  $J_0$  observed by POES satellites over 6–9 MLT for >30 keV. (h and i) The same as Figures 1f and 1g but for electron energy of >100 keV. (j and k) The same as Figures 1f and 1g but for electron energy of >300 keV.

including the electron energy spectrum, plasma density, wave normal distribution, wave frequency spectrum, and the latitudinal dependence of the wave normal angle distribution and plasma density etc. In the present paper, we obtained most of these required parameters from the Van Allen Probe observations near the equator and only made assumptions on the latitudinal dependence of the wave normal distributions and plasma density.

### 3. Results of the Conjunction Event

We conduct a quantitative analysis for hiss-driven electron precipitation for an event, which occurred on 8 November 2012. During this time period, the geomagnetic activity was quiet overall, with  $Kp$  less than 2 (not

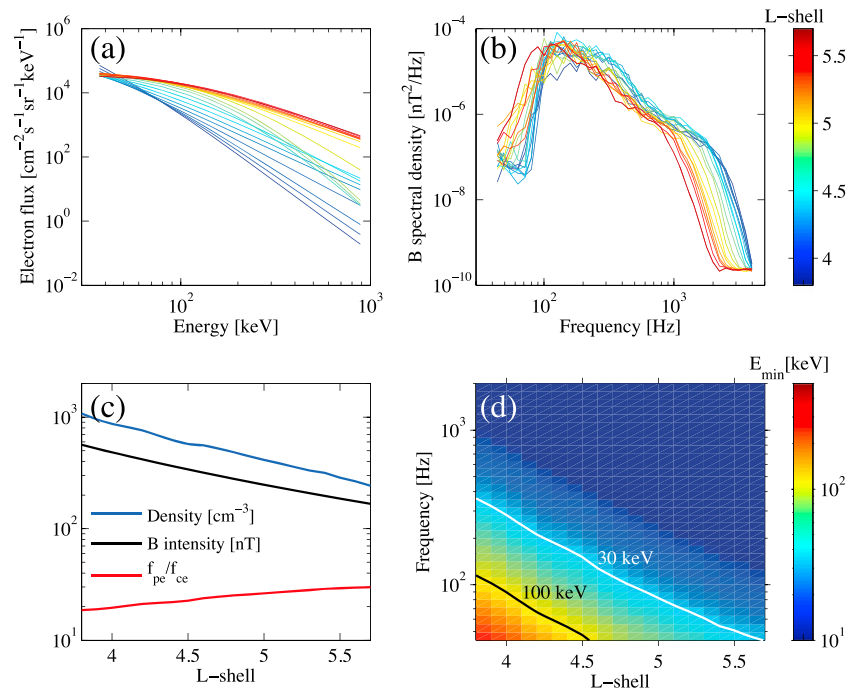
shown), but there was a modest disturbance in  $AL$  (Figure 1a) down to  $-300$  nT between 10 and 13 UT, followed by a quiet period during the rest of the day (except for a small disturbance in  $AL$  at around 22 UT). Figure 1 shows Van Allen Probe A observations of waves and electron flux near the dawn sector in the equatorial magnetosphere and POES observations of  $J_{90}$  and  $J_0$  at various energies ( $>30$  keV,  $>100$  keV, and  $>300$  keV) in a similar range of MLT (6–9). Here the electron fluxes measured by POES satellites, which passed over 6–9 MLT, were averaged into  $0.5 L \times 1$  h (in UT) bins. During Orbit 186, as shown from the upper hybrid line in Figure 1b, the plasmopause was located at  $\sim 4 R_E$  near the dawn sector, and weak chorus waves were observed outside the plasmopause (Figures 1c and 1d).  $J_0$  was very weak during the time interval of Orbit 186, as shown at various energies in Figures 1g, 1i, and 1k associated with these weak chorus waves, although the measured  $J_{90}$  was sufficiently above the noise level. During Orbit 187, the plasmopause location moved outside the apogee of Probe A ( $>\sim 5.8 R_E$ ), and strong hiss waves were observed throughout the outer plasmasphere in association with the modest disturbance in  $AL$ . Interestingly, a pronounced electron precipitation at  $>30$  keV (Figure 1g) and  $>100$  keV (Figure 1i) was also observed during the period (11–17 UT) when Van Allen Probe A observed strong hiss. Note that  $J_0$  at  $>300$  keV (Figure 1k) was close to the noise level and thus is excluded from further quantitative analysis. During Orbit 188, no waves were observed outbound (18–21 UT) in the postmidnight sector, followed by modest whistler mode waves observed during 22–24 UT near the dawn sector. These modest whistler mode waves were associated with modest electron precipitation at  $>30$  keV and  $>100$  keV observed by POES during the same time interval (22–24 UT). Note that Van Allen Probe B observed very similar features, since the two satellites were very close to each other on this day. Overall, low-altitude electron precipitation at  $>30$  keV and  $>100$  keV exhibited remarkable correlation with whistler mode wave intensity observed near the magnetic equator.

In the present paper, we focus on the period of 13–16 UT (indicated by the red horizontal arrow in Figure 1), when Van Allen Probe A traveled from the apogee ( $\sim 5.8 R_E$ ) to lower  $L$  shells down to 3.8, and perform a quantitative analysis of electron scattering driven by plasmaspheric hiss. As shown in equation (1), calculation of  $R$  requires the electron energy spectrum ( $J(E)$ ), which can be obtained from the in situ MagEIS measurements in the equatorial magnetosphere. Based on the MagEIS electron flux observed during the analyzed period (13–16 UT), we fit the energy spectrum using a relativistic Kappa-type distribution [Xiao *et al.*, 2008] at each  $L$  shell as shown in Figure 2a (color coded for various  $L$  shells). Note that this Kappa-type distribution obeys a power law distribution at the relativistic energies [e.g., Lu *et al.*, 2011]. The energy spectrum was softer at lower  $L$  shells and became harder at larger  $L$  shells. Although the wave spectrum (Figures 1c and 1d) did not show significant variation during 13–16 UT, the wave frequency corresponding to the peak wave power became slightly lower at higher  $L$  shells (Figure 2b). The total plasma density calculated from the upper hybrid line measured on EMFISIS and a dipole magnetic field model are used to calculate the ratio of the plasma to electron cyclotron frequency ( $f_{pe}/f_{ce}$ ) and the electron minimum cyclotron resonant energy ( $E_{min}$ ) at the geomagnetic equator, as shown in Figures 2c and 2d. We confirmed that a dipole magnetic field model agrees well with the observed local magnetic field measurements during this time interval associated with modest geomagnetic activity near the dawn sector.  $f_{pe}/f_{ce}$  increases with increasing  $L$  shells over 3.8–5.8 inside the plasmasphere.  $E_{min}$  decreases with increasing  $L$  shells and frequencies (Figure 2d) and quickly increases with increasing magnetic latitudes [Bortnik *et al.*, 2011]. Hiss with wave frequencies over 100–200 Hz, corresponding to the peak wave power, resonates with electrons with  $E > \sim 30$  keV at  $L \sim 2$  and  $E > \sim 100$  keV at  $L \sim 4$ .

We use the measured wave frequency spectrum and plasma density by Van Allen Probe A to calculate bounce-averaged electron pitch angle diffusion coefficients at the loss cone at various  $L$  shells and energies (Figure 3). Since hiss was observed continuously over 13–16 UT (Figures 1c and 1d), we assume that hiss was always present during the interval of the conjunction event in a region with the  $L$  shell of 3.8–5.8 over 6–8.5 MLT. We also assume a constant hiss spectrum distributed up to  $45^\circ$  of the magnetic latitude ( $\lambda$ ) following the ray-tracing results [e.g., Bortnik *et al.*, 2011]. Total plasma density along the magnetic field line is assumed to vary following

$$n_e(\lambda) = n_{e0} / \cos(\lambda)^{2a}, \quad (2)$$

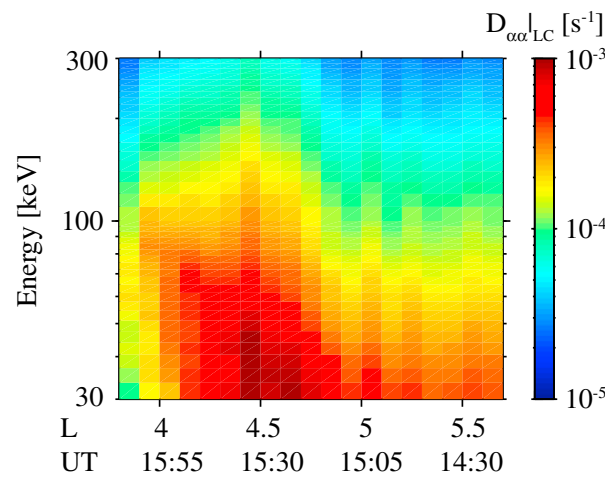
where  $n_e$  is the local plasma density at magnetic latitude  $\lambda$  and  $n_{e0}$  is the equatorial value. The parameter  $a$  is



**Figure 2.** Parameters used to calculate electron pitch angle diffusion coefficients at various  $L$  shells. (a) Electron energy spectrum obtained by fitting a relativistic Kappa-type function based on MAGEIS electron flux measurements during 13–16 UT on 8 November 2012, when Van Allen Probe A traveled through  $L$  shells from 5.8 to 3.8, color coded for various  $L$  shells. (b) Magnetic field spectral density observed by the EMFISIS instrument on Van Allen Probe A, color coded for various  $L$  shells. (c) Total plasma density (blue) calculated from the upper hybrid line, background magnetic field (black), and the ratio of plasma frequency to electron cyclotron frequency ( $f_{pe}/f_{ce}$ ) (red) as a function of  $L$  shell. (d) Electron minimum cyclotron resonant energy at the geomagnetic equator as a function of  $L$  shell and wave frequency, where the white and black lines represent the  $E_{min}$  contours of 30 and 100 keV.

set to 1 according to the statistical value obtained inside the plasmaphere [e.g., Denton et al., 2006]. The wave normal distribution of hiss is assumed to be Gaussian and given by

$$g(\theta) = \exp \left[ - \left( \frac{\tan \theta - \tan \theta_m}{\tan \theta_w} \right)^2 \right] (\theta_{lc} \leq \theta \leq \theta_{uc}), \quad (3)$$



**Figure 3.** Bounce-averaged electron pitch angle diffusion coefficients at the bounce loss cone as a function of  $L$  shell (and UT on 8 November 2012 when Van Allen Probe A traveled through the corresponding  $L$  shell) and electron energy.

where  $\theta$  is the wave normal angle,  $\theta_m$  is the peak,  $\theta_w$  is the angular width, and  $\theta_{lc}$  and  $\theta_{uc}$  are the lower and upper cutoffs to the wave normal distribution, outside which the wave power is zero. The adopted wave normal angle distribution is shown in Table 1, which varies with magnetic latitudes, consistent with ray-tracing results [Bortnik et al., 2011] and statistical Cluster observations [Agapitov et al., 2013]. Figure 3 shows bounce-averaged pitch angle diffusion coefficients at the loss cone, which is strongly energy-dependent with larger values at lower energies at most  $L$  shells except at  $L < \sim 4.2$ .  $D_{\alpha\alpha}(E)|_{\alpha=\alpha_0}$  reaches its maximum at  $L \sim 4.4$  for all electron energies corresponding to where hiss amplitudes are observed to be strongest up to 115 pT (blue dashed line in Figure 4b).

**Table 1.** Wave Normal Distributions of Plasmaspheric Hiss at Various Magnetic Latitudes

$ \lambda $	$\theta_m$	$\theta_{IC}$	$\theta_{UC}$	$\theta_w$
0° – 5°	0°	0°	10°	5°
5° – 10°	10°	0°	20°	10°
10° – 15°	20°	0°	30°	15°
15° – 20°	25°	10°	40°	20°
20° – 25°	30°	20°	50°	25°
25° – 30°	35°	30°	60°	30°
30° – 35°	45°	40°	65°	35°
35° – 40°	55°	50°	70°	40°
40° – 45°	65°	60°	75°	50°

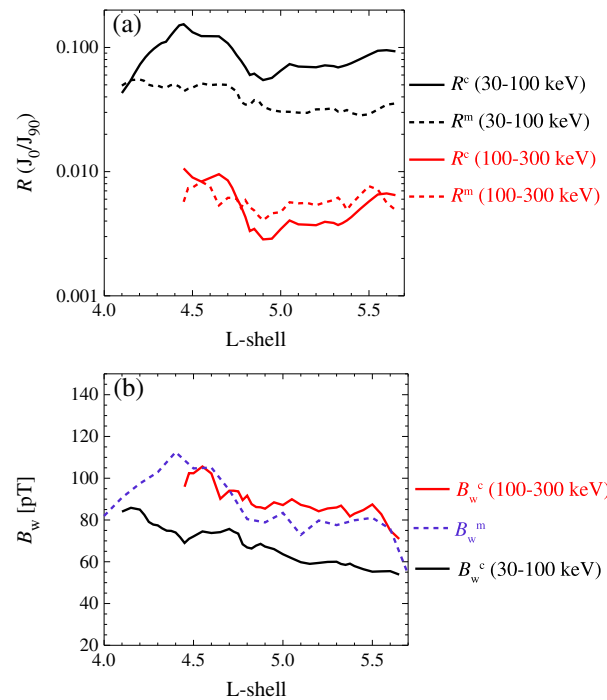
As discussed above, the wave frequency spectrum showed insignificant variation between 13 and 16 UT, when Van Allen Probe A traveled from 5.8 to 3.8  $R_E$  and from 6 to 8.5 MLT. Between 15:03 and 15:05 UT, the NOAA 17 satellite traveled through the region with  $L$  shells from  $\sim 3.8$  to  $\sim 6$  near 7.5–8 MLT at the low altitude of  $\sim 800$ –850 km. Therefore, it is reasonable to assume that the precipitated electron fluxes measured by NOAA 17 at  $L$  shells over 3.8–5.8 in the postdawn sector are predominantly due to

pitch angle scattering by hiss in the conjugate equatorial magnetosphere.

We calculated the ratio ( $R^c$ ) in the energy channel of 30–100 keV and 100–300 keV using equation (1) based on the parameters shown in Figure 2 and pitch angle diffusion coefficients shown in Figure 3. We also obtained measured ratio ( $R^m$ ) of  $J_0/J_{90}$  for the energy range of 30–100 keV and 100–300 keV by NOAA 17. Note that the electron fluxes measured over 30–100 keV (100–300 keV) were obtained by subtracting  $>100$  keV ( $>300$  keV) electron fluxes from the values of  $>30$  keV ( $>100$  keV). Comparison of the calculated and directly measured  $R$  (Figure 4a) exhibits fairly good agreement particularly for 100–300 keV. Note that  $R$  (Figure 4a) and  $B_w^c$  (Figure 4b) values are only shown for the  $L$  shells, where electron counts measured by 0° and 90° telescopes are sufficiently above the background level ( $>5$  s $^{-1}$ ). This is why no values are shown at  $L < 4.4$  in the 100–300 keV

channel in Figures 4a and 4b.  $R^c$  at 30–100 keV is larger than  $R^m$ , but the difference is generally within a factor of about 2.  $R$  is generally larger at  $L$  shells of 4.3–4.7 compared to that at larger  $L$  shells for both measured and calculated values.

Figure 4b shows calculated hiss wave amplitudes from the measured  $R$  for two energy channels of 30–100 keV and 100–300 keV separately, using the method described in section 2.3. The calculated hiss wave amplitudes ( $B_w^c$  in solid red line) from  $R^c$  (100–300 keV) show remarkable agreement with the directly measured hiss wave amplitudes from EMFISIS ( $B_w^m$  in dashed blue line). Note that although the measured  $B_w^m$  ( $>80$  pT) is larger than the hiss wave amplitudes of 10 pT, which was used to represent the long-term averaged value in the study of *Abel and Thorne [1998]*, it is comparable to the hiss wave amplitudes observed by the CRRES during active times [*Meredith et al., 2007*]. Hiss wave amplitudes ( $B_w^c$  in solid black line) calculated from  $R^c$  (30–100 keV) are slightly lower than the measured values at  $L > 4.7$ , while the consistency becomes worse over the  $L$  shells of 4.0–4.7. Nevertheless, the difference in measured and calculated wave amplitudes was within a factor of 2, which suggests that measured values of  $R$  can be used to provide a realistic estimate of the hiss wave



**Figure 4.** Comparison of measured and calculated  $R$  and hiss wave amplitude ( $B_w$ ). (a)  $R (J_0/J_{90})$  calculated using measured wave spectral intensity and plasma parameters ( $R^c$ : solid black line) and directly measured from two-directional POES measurements ( $R^m$ : dashed black line) for 30–100 keV electrons. Red lines represent similar quantities to black lines but for electron energies of 100–300 keV. (b) Directly measured hiss wave amplitudes on EMFISIS integrated over 40–4000 Hz ( $B_w^m$ : blue dashed line) and calculated hiss wave amplitudes using  $R^m$  of 30–100 keV ( $B_w^c$ : black line). Red line is similar to the black line but for electron energy of 100–300 keV.



amplitudes at various  $L$  shells. Note that since  $\sim 30$  keV is very close to the electron minimum resonant energy, any change in real parameters (such as the latitudinal dependence of the wave normal angle distribution and plasma density) observed in space compared to our assumptions could bring a substantial portion of the  $\sim 30$  keV electron populations out of resonance. This will cause larger changes in precipitated electron fluxes and  $R$  in the 30–100 keV energy channel, and thus,  $R$  (30–100 keV) provides a larger uncertainty in estimating hiss wave amplitudes compared to the 100–300 keV channel.  $R^m$  (100–300 keV) is likely to be more robust compared to  $R^m$  (30–100 keV) to calculate hiss wave amplitudes, since 100–300 keV electrons almost always resonate with plasmaspheric hiss over its typical frequency range.

Note that in this conjunction event analysis, parameters near the equator were taken based on the in situ Van Allen Probe A observation. However, in the off-equatorial region, we assume the latitudinal dependence of the wave normal angle distribution and plasma density based on the statistical observations [Agapitov *et al.*, 2013; Denton *et al.*, 2006]. We also tested different wave normal distributions (e.g., quasi-parallel wave normal angle distributions along the field line) and different values of  $a$  (in equation (2)) over 0–1 to model different density distributions along the field line, and the results (not shown) are very similar to those shown in Figure 4. This also indicates that this technique is relatively insensitive to these assumed parameter changes and provides robust and reliable estimates of hiss wave amplitudes.

#### 4. Summary and Discussion

Using a physics-based technique, we quantify hiss-driven electron precipitation during a conjunction event, when Van Allen Probe A observed plasmaspheric hiss in the equatorial magnetosphere and NOAA 17 measured electron precipitation at the conjugate low altitudes. We calculated the ratio of electron fluxes ( $R$ ), which could be measured by two-directional POES particle detectors, based on observed wave and plasma parameters. Comparison between the measured and calculated  $R$  over the energy of 30–100 keV and 100–300 keV showed fairly good agreement at various  $L$  shells. The measured  $R$ , in turn, was used to calculate the hiss wave amplitudes, which showed remarkable agreement with the directly measured hiss wave intensity particularly at 100–300 keV. Note that the 100–300 keV energy channel is more robust and reliable to infer hiss wave amplitudes compared to the energy channel of 30–100 keV. This suggests that this technique, based on quasi-linear diffusion theory, can be used to estimate hiss wave amplitudes from the low-altitude energetic electron measurements and thereby construct the spatiotemporal evolution of hiss wave intensity on a global scale using observations from multiple POES satellites. Such a data-driven dynamic hiss wave distribution could be critically important for the evaluation of electron loss from the inner magnetosphere over a broad energy range (tens of keV to several MeV).

This technique will only be reliable in regions where measured electron fluxes by POES satellites are sufficiently above the noise level, which limits its use to a certain range of  $L$  shells, i.e., at  $> \sim 4 R_E$  in the event studied here. However, ray-tracing simulations of the propagation of hiss [Bortnik *et al.*, 2011; Chen *et al.*, 2012] provide insightful information on how hiss wave power scales over various  $L$  shell ranges. The combination of available hiss wave amplitudes inferred from low-altitude electron precipitation, together with in situ equatorial wave observations and ray-tracing simulations, has the potential to provide more complete information on the spatiotemporal evolution of hiss in the entire plasmasphere during various geomagnetic activities, but this will require further development.

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