## Recombination-line intensities for hydrogenic ions - I. <br> Case $B$ calculations for $\mathrm{H}_{\mathrm{I}}$ and $\mathrm{He}_{\text {II }}$

 Summary. The relative intensities of $\mathrm{H}_{\mathrm{I}}$ and $\mathrm{He}_{\text {II }}$ recombination lines arecalculated, including full collisional effects, for a considerably larger range of
temperature, density and principal quantum numbers than before. Case B of
Baker \& Menzel is assumed, and tables of line opacities are also presented to
enable the assumption of negligible optical depth in all but the Lyman series to be
checked. Collisional excitation of the $n=3$ levels from both $n=1$ and $n=2$ states is
considered, and is found to invalidate Case B theory in some conditions which
depend on electron density and the Lyman- $\alpha$ escape probability. The regimes of
temperature and particle density for which Case B is valid are discussed. Newly
calculated collision strengths for the $n=1,2$ and 3 states of $\mathrm{He}^{+}$are tabulated.

## 1 Introduction

With the increasing sensitivity and resolution available to infrared spectroscopy, recombination lines can be observed in many types of astronomical objects at ever-increasing wavelengths. In order to interpret these observations, which can be uniquely useful because of the relative simplicity of the physical process controlling the formation of infrared recombination lines, antum necessary to have theoretical line intensities for transitions involving levels with numbers larger than previously considered. Moreover, in many of these objects, such as




 Menzel (1938) for which the Lyman lines are taken to be optically thick. In the Case B *Staff Member, Quantum Physics Division, National Bureau of Standards. Permanent address: Joint Institute for Laboratory Astrophysics, National Bureau of Standards and University of
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2 Formulation
In the determination of the level populations, $N_{n l}$, of hydrogenic systems, we follow the approach of Brocklehurst $(1970,1971)$, and make use of the computer program used in the latter work with a number of modifications and improvements. It is convenient to introduce the departure coefficients $b_{n l}$ defined in terms of the Saha-Boltzmann populations by
(1) parameterized by the Lyman alpha escape probability under which these effects can be neglected. The resulting relative intensities and opacity factors, which are tabulated in full on microfiche, are discussed in Section 5. collisional transitions from levels with $n=1$ and $n=2$ to higher levels, and derive conditions
 approxim can be easily calculated.
In a standard Case B calculation, the effect of collisional excitations from $n=1$ and $n=2$ states is ignored. The range of densities of interest for Wolf-Rayet stars includes a regime in which collisional excitation from the $n=2$ states may become important. In this paper we restrict ourselves to densities for which the standard Case B assumptions are valid. A subsequent
The recombination of a high-density plasma has been treated previously by a number of authors. The connection between the so-called collisional-radiative recombination theory of Bates, Kingston \& McWhirter (1962) and the approach used by Brocklehurst (1970, 1971) and by ourselves is shown by Burgess \& Summers (1976) and Summers (1977). Our Case B theory is identical to the 'depopulated $n=2$ ' case treated by Summers (1977), apart from numerical results are essentially identical to those of Summers (1977) in the few cases where comparison can be made, although the goals of his paper are very different from ours.
 upper and lower principal quantum numbers up to 50 and 29 , respectively. For Hi we cover the temperature and density ranges $3 \leq \log T_{\mathrm{e}} \leq 4.5$ and $2 \leq \log N_{\mathrm{e}} \leq 10$, while for HeII, $3.5 \leq$
 (1971). In Section 2 we give a brief description of the method of calculation and outline the most important differences from the work of Brocklehurst. Section 3 contains a detailed


## D. G. Hummer and P. J. Storey



 proportional to $(2 l+1)$, so that $b_{n l}=b_{n}$. A matrix condensation method is used to reduce the



 stage are used to start an iterative solution for the $b_{n l}$. Brocklehurst (1971) defined a value of
Recombination-line intensities for hydrogenic ions 803






 energy-changing collisional processes. There are also some differences in the numerical methods used to evaluate the infinite sums of radiative rates. In the calculation of $b_{n}$, Brocklehurst
 transitions were included in which the principal quantum number changed by $\pm 1$. Although transitions with $|\Delta n|=1$ dominate in the total rate for energy-changing collisions, this inconsistency in the treatment of collisions in the two stages of the calculation led to difficulties in matching the values of $b_{n l}$ and $b_{n}$ at $n=n_{\mathrm{c}}$ (Brocklehurst 1971). These difficulties become worse at densities larger than those considered by Brocklehurst. In the calculations described here, all electron-induced energy-changing collisional processes are incorporated, both in the calculation of $b_{n}$ and the subsequent calculation of $b_{n l}$. In addition, the same cross-sections and the same numerical methods are used in both stages.

## 3 Atomic rate coefficients and numerical methods

### 3.1 CALCULATION OF $b_{n}$

We describe the atomic processes included in the calculation of $b_{n}$ and compare our methods with those of Brocklehurst (1970):

### 3.1.1 Bound-bound radiative transitions <br> As in Brocklehurst (1970)

### 3.1.2 Radiative recombination

As in Brocklehurst (1970)
Brocklehurst (1970) used a mixture of impact-parameter, correspondence principle, and binary encounter cross-sections for these transitions. Here, the fits to various classical and semi-classical results of Percival \& Richards (1978) are used. A more complete background to these cross-sections is given by Percival \& Richards (1975)
3.1.4 Collisional ionization and three-body recombination
We use the symmetrized binary encounter ionization cross-section of Burgess \& Percival (1968, equation 37). Burgess \& Summers (1976) have added a contribution allowing for distant
 Using the formulae of Summers (1979), we find that the maximum contribution to the collision rate from this logarithmic, quantal part of the cross-sections is 4 per cent for the $n=3$ state of $\mathrm{He}^{+}$ at 100000 K . The quantal contribution declines rapidly as $n$ increases and electron temperature

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804 D.G.Hummer and P.J.Storey
decreases. We therefore neglect this term in the cross-section. Binary-encounter ionization
cross-sections are also discussed by Percival \& Richards (1975), who introduce a modification to
the binary encounter result for H at low energies. We exclude this modification since it is only of
importance at energies, and for states, for which collisional ionization is an unimportant process
in the conditions considered here. Brocklehurst (1970) does not specify how he calculated the
rates for these processes, but it seems likely that he also used a binary encounter method, which is
suitable for large energy transfers. Rate coefficients for three-body recombination are obtained
from those for collisional ionization using the principle of detailed balance.
We employ the matrix condensation technique first applied to this problem by Burgess \& Summers (1969). A description of the method is given by Brocklehurst (1970) and by Burgess \& Summers (1976)

### 3.1.5 Numerical methods

### 3.2.1 Bound-bound radiative transitions

Brocklehurst (1971) evaluates transition probabilities directly from the expression of Gordon (1929) in terms of the hypergeometric functions ${ }_{2} \mathrm{~F}_{1}$. He calculates these functions directly from series expansions. We use the same method for low principal quantum numbers ( $n \leq 200$ ), but for higher values, which are required in our modified treatment of the cascade problem, we use
recursion relations to generate the functions in order to avoid cancellation difficulties (see for recursion relations to generate the functions in order to avoid cancellation difficulties (see for
example, van Regemorter et al. 1979).

### 3.2.2 Radiative recombination

As in Brocklehurst (1971).
3.2.3 l-changing collisions
As in Brocklehurst (1971), we use the semi-classical impact-parameter treatment for degenerate
 region of incident particle impact parameter, $r_{p}$, are defined. The boundary of these two regions, $r_{\mathrm{p}}=r_{\mathrm{s}}$, is defined by

## $\sum_{l^{\prime}=l \pm 1} P_{n l, n l^{\prime}}=1 / 2$,

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where $P_{n l, n l^{\prime}}\left(E, r_{\mathrm{p}}\right)$ is the probability of a transition from $n l \rightarrow n l^{\prime}$ at an incident energy $E$, and impact parameter $r_{p}$. The cross-section for collisions out of $n l$ is then given by
(3)
where it has been assumed that in the strong coupling region the average probability of a
transition is $1 / 2$, and $r_{u}$ is an upper cut-off, discussed by Brocklehurst (1971). The probability s، $\downarrow$ S.n






 cross-section is proportional to the oscillator strength, so that for all impact parameters,
( $t$ )
The $C_{n}$ obtained with this assumption differ by at most a few per cent from those obtained by Brocklehurst's method, but vary smoothly with $l$. In practice, $C_{n l, n l+1}$ is calculated and $C_{n l+1, n l}$ is obtained from detailed balance considerations. In the calculation of the rate coefficients for these collisions we have included the effects of electrons, protons and $\mathrm{He}^{+}$ions. The assumed abundances are discussed in Section 4.

### 3.2.4 Energy-changing collisions

We use the cross-sections described in Section 3.1.3 above. As the collision rates for $\Delta l=l^{\prime}$ $-l= \pm 1$ dominate, we consider only those transitions. Because the collision rates among the nearly degenerate $l$-states for a given value of $n$ are always much larger than for $n$-changing
collisions, the precise $l$-dependence of the $n l \rightarrow n^{\prime} l \pm 1$ cross-sections is not of primary importance. Moreover, the dominant contribution comes from transitions with $|\Delta n|=1$.


## $\Delta l= \pm 1$

$$
\Delta l \neq 1
$$

## $Q_{n l, n \pm 1 l^{\prime}}(E)=\frac{\int_{n l, n \pm 1 l \pm 1}}{f} Q_{n, n \pm 1}(E)$,

$$
=0
$$(5)

 Bethe approximation for collisional excitation. For transitions with $|\Delta n|>1$ we use the simpler approximation,
$Q_{n l, n^{\prime} l+\Delta l}(E)=Q_{n n^{\prime}}(E)$, $=0$,

## $\Delta l \neq \pm 1$.

These assumptions maintain the correct total number of $n \rightarrow n^{\prime}$ transitions, and rely on the very rapid $l$-changing collisions with $\Delta n=0$ to correctly distribute population among the $l$-substates









3.2.6 Numerical methods
As outlined in Section 2, the iterative method of evaluation of $b_{n l}$ is similar to that used by Brocklehurst (1971). The differences lie in the treatment of the radiative cascade and collisional contributions to the populations of each sublevel $n l$. In Brocklehurst's treatment it was assumed
 $n>n_{\mathrm{c}}$. To obtain the total radiative cascade rate to a given level, $n l$, Brocklehurst sums the first
 transition probability from some higher state $n^{\prime} l^{\prime}$ to $n l$. This quantity is slowly varying as a function of $1 / n^{\prime 2}$. The fit is then integrated analytically. We replace this latter procedure with five-point Gaussian integration in the variable $1 / n^{\prime 2}$. These changes do not represent a great improvement in accuracy over Brocklehurst's approach, but arise from our desire to use, as far as is possible, the same techniques in the first-stage calculation of $b_{n}$ as in the subsequent calculation
 terms and carrying out a numerical integration for the remainder. In this case a 16-point Gaussian integration in the variable $1 / n^{2}$ is used. Again the same method is used in both stages of the calculation.

## 4 Collisional excitation of higher levels from $n=1$ and $n=2$

The 'Case B' theory used here is based on the assumption that the level populations for $n \geq 3$ are

 collisional transitions from $n=1$ and 2 may make an appreciable contribution to the populations
 be derived under which these collisional transitions are negligible and that the users of our tables take these limits seriously.
 collisional transitions from all levels with larger principal quantum numbers is given in terms of an effective recombination rate constant $\bar{\alpha}_{n l}$, which is defined by
$N_{\mathrm{e}} N_{+} \bar{\alpha}_{n l}=N_{\mathrm{e}} N_{+}\left(\alpha_{n l}+N_{\mathrm{e}} \beta_{n l}\right)+\sum \sum N_{n^{\prime} l^{\prime}}\left(A_{n^{\prime} l^{\prime}, n l}+N_{\mathrm{e}} C_{n^{\prime} l^{\prime}, n l}\right)$,
(8)
where $\alpha_{n l}$ and $\beta_{n l}$ are the direct radiative and three-body recombination rate coefficients, $A_{n^{\prime} l^{\prime}, n l}$
is the Einstein coefficient and $C_{n^{\prime} l^{\prime}, n l}$ is the collisional rate coefficient for the $\left(n^{\prime}, l^{\prime}\right) \rightarrow$
( $n, l$ ) transition. From Tables 1 and 2 , which give values of $\bar{\alpha}_{n l}$ for $n=2$ and 3 in hydrogen and
ionized helium, respectively, at selected values of $N_{\mathrm{e}}$ and $T_{\mathrm{e}}$, it is clear that these coefficients
are smoothly decreasing functions of temperature and are almost independent of density. Rate
coefficients for electron-induced transitions between levels with $n=1,2$ and 3 are required for
both H and $\mathrm{He}^{+}$. We discuss these rate coefficients in the Appendix. The Appendix contains the

[^0]$$
\stackrel{\delta}{\infty}
$$


## $4.11 \rightarrow 2$ AND $1 \rightarrow 3$ COLLISIONAL TRANSITIONS

results of a new close-coupling and distorted-wave calculation of collision strengths for $\mathrm{He}^{+}+e$, which provides rates for the $1 \rightarrow 3$ and $2 \rightarrow 3$ transitions substantially more accurate than those currently in the literature.
The ground-state population is determined by the balance of ionizing and recombining mechanisms, and consequently its value in any particular circumstance cannot be obtained from recombination theory alone. However, we can derive a limiting value for the degree of ionization which ensures that collisional transitions from the ground state are negligible. The population of level ( $n, l$ ) will be unaffected by collisions from $n=1$ if

## $N_{1 s} N_{\mathrm{e}} C_{1 s, n l} \ll N_{\mathrm{e}} N_{+} \bar{\alpha}_{n l}$.

(9)
We define a critical degree of ionization $\left(N_{1 s} / N_{+}\right)_{n l}^{\text {crit }}$, such that the total rate of collisional population of the $(n, l)$ level is a small fraction, $p$, of the population rate from higher states, by (10)


 then the tabulated Case B results will still be valid providing collisional $2 \rightarrow 3$ transitions are negligible; this issue is discussed in the next Section 4.2. On the other hand, if collisional excitation of $n=3$ from the ground state is significant, then $1 s \rightarrow n l$ collisions are probably contaminating level populations for all levels, and the results of this paper are simply

$$
\begin{aligned}
& \text { Table 3. Critical values of } N(1 s) / N_{+} \text {such that } 1 s \rightarrow n l \text { collisional excitation is } \\
& 10 \text { per cent of the total population rate for } \mathrm{H} \text { and } \mathrm{He}^{+} \text {. }
\end{aligned}
$$ inapplicable.

$$
\mathrm{H}_{\mathrm{I}}\left(N_{\mathrm{e}}=10^{6} \mathrm{~cm}^{-3}\right)
$$

$\square$

$$
\cdots \stackrel{7}{1}
$$

$4.22 \rightarrow 3$ COLLISIÓNAL TRANSITION
 can be determined from Case B theory by a simple extension, at least if collisional $1 \rightarrow 2$ transitions are negligible. The statistical equilibrium equations for the $2 s$ and $2 p$ levels, respectively, can be written as
$N_{+} \bar{\alpha}_{2 s}+N_{2 p} C_{2 p, 2 s}+N_{1 s} C_{1 s, 2 s}=N_{2 s}\left[A_{2 q} N_{\mathrm{e}}^{-1}+C_{2 s, 1 s}+C_{2 s, 2 p}+C(2 s)\right]$



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| :---: | :---: |

 at the highest densities it is not strongly temperature-dependent. Simple approximate forms



## (e8I) (981) $\beta_{\mathrm{c}}^{(3)}=N_{\mathrm{e}}\left(3 C_{p}+C_{s}\right) / 3 A_{2 p, 15}, \quad 3 N_{\mathrm{c}} C \gtrdot N_{\mathrm{e}} \max \left(C_{s}, C_{p}\right) \gg A_{2 q}$. (18c)  $2 \times 10^{-13} N_{\mathrm{e}} Z^{-4}$. In regions (1) and (2) the dominant mechanism for depopulating the $2 p$ level is collisional transitions to $2 s$ followed by two-quantum decay, while in region (3), it is collisions to higher levels.

## $D=\left[R_{s}\left(C+C_{p}\right)+3 C C_{p}\right]\left[1+\left(\beta / \beta_{c}\right)\right]$.

(19)
The only other occurrence of $\beta$ in equations (16) is in the combination $R_{p}+C$, which can be written as
Thus the populations of the $n=2$ levels are manifestly independent of $\beta$ when $\beta \ll \beta_{c}$; thus one can

 contrast, for lower densities the $\beta=1$ limit is probably appropriate, as the condition $\beta \ll \beta_{\mathrm{c}}$ is likely


 respectively. Hence Case B is relatively easy to achieve for the higher Lyman lines, but not for Lyman $\alpha$.

## For $2 \rightarrow 3$ collisions to be negligible, we require

(Iz)
Let us first assume that the terms in equations (16) containing $N_{1 s}$ are negligibe, according
 $n=3$ levels such that the total rate of collisional population of the $n=3$ states from $n=2$ is a small fraction, $p$, of the population rate from all higher states. That is

[^1]then equations (16) take the form
$N_{2 l}=(2 l+1) N_{\mathrm{e}} N_{+} \bar{\alpha}_{2} /\left(A_{2 q}+3 A_{2 p, 1 s} \beta+\gamma N_{\mathrm{e}}\right)$,
813 Recombination－line intensities for hydrogenic ions

| $\begin{gathered} \frac{10}{\tilde{0}} \\ \frac{0}{6} \\ \frac{0}{8} \end{gathered}$ | $\stackrel{+}{+} \stackrel{+}{\sim_{2}^{\prime}}$ | $0 \stackrel{O}{0}_{0}^{\infty}$ |
| :---: | :---: | :---: |
|  | $\stackrel{+}{\mathbf{O}}$ | $\begin{aligned} & 8 \\ & 8 \\ & 8 \\ & \hline \end{aligned}$ |
|  | $\begin{aligned} & \text { n } \\ & 0 . \\ & 0.0 \\ & 0 . \end{aligned}$ |  |
|  | $\begin{gathered} \text { E } \\ \text { E } \\ 0 \\ \text { D } \end{gathered}$ | 㔛 垔 |

values of the Lyman $\alpha$ escape probabilities $\beta \gg \beta_{\mathrm{c}}$ ，where $\beta_{\mathrm{c}}$ is given by Figs 1 and 2．The species and relative number of ions causing $l$－changing collisional transitions，are given in Table 5．We assume throughout that $N(\mathrm{He}) / N(\mathrm{H})=0.1$ ．As He II recombination lines formed under rather different sets of conditions are observed in different types of objects，such as planetary nebulae and Wolf－Rayet stars，we have made trial calculations for He II which reflect the ionic abundances in these two classes of object．The＇PN＇calculations were carried out with the abundances given in Table 5．The＇WR＇calculation assumes that the plasma is composed entirely
of $\mathrm{He}^{2+}$ and electrons，so that $N\left(\mathrm{He}^{2+}\right)=N(\mathrm{He})=0.5 N_{\mathrm{e}}$ ．Comparing the PN and WR results at
 about 12 per cent，which occurs in comparing lines in the infrared with lines in the ultraviolet．

 differences，we tabulate results only for the PN model．Results are given for all transitions
between pairs of upper and lower principal quantum numbers $n_{\mathrm{u}}$ and $n_{l}$ ，respectively，in the For each ion and each pair of parameters $\left(T_{\mathrm{e}}, N_{\mathrm{e}}\right)$ we give both the relative intensities and the
opacity factors，as defined below．In Tables 6 and 7 we reproduce the tabular material from the fiche for the case of H at $N_{\mathrm{e}}=10^{4} \mathrm{~cm}^{-3}$ and $T_{\mathrm{e}}=10^{4} \mathrm{~K}$ ．At the top of each table appears the identification of the spectrum，the values of $T_{\mathrm{e}}$ and $N_{\mathrm{e}}$ and of the parameter NC．The second line
 recombination coefficient for the reference transition 4－3 or 4－2 as indicated，the power in the reference transition in $\mathrm{erg} \mathrm{cm}^{-3} \mathrm{~s}^{-1}$ ，and finally the coefficients $\bar{\alpha}_{2 s}$ and $\bar{\alpha}_{2 p}$ defined by equation （8）．The main body of the upper part of the table gives the energy in each transition $n_{\mathrm{u}}$ to $n_{l}$ relative to unit energy in a reference transition，which is $4-2$ for HI and $4-3$ for $\mathrm{He}^{+}$．
wer part of the table contains，for each transition，the opacity factor $\Omega_{n, n^{\prime}}$ defined as


 ing to a path length of $L \mathrm{~cm}$ is then

$$
\tau_{n n^{\prime}}=N_{\mathrm{e}} N_{+} \Omega_{n n^{\prime}} L
$$

Note that for many transitions $\Omega_{n n^{\prime}}<0$ ，which indicates population inversions and the potential for maser action in those transitions．

The theory used in these calculations assumes that $\left|\tau_{n n^{\prime}}\right| \ll 1$ for $n^{\prime}>n>1$ ．When using the tabulated line intensities for an $n^{\prime} \rightarrow n$ transition in a particular astrophysical context，it is necessary to confirm that the optical depth given by equation（30）is small in all transitions， especially those involving states below $n, n^{\prime}$ ．In this paper we do not calculate populations for
 the estimates of $2 s$ and $2 p$ level populations given by equations（16），these quantities can be calculated from equation（29）．The stimulated emission factor can be ignored．
D．G．Hummer and P．J．Storey

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| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  |  |  |  |  |  |  |  |  |  |  | 50－ヶ¢＇z | S0－G¢ $\frac{1}{\text { ¢ }}$ | 1 |
|  |  |  |  |  |  |  |  |  |  |  | 50－18．1 | 50－2E．2 | C0－76．${ }^{\text {c }}$ | 61 |
|  |  |  |  |  |  |  |  |  |  | ¢0－2¢．t | S0－12－1 | ¢0－60＇2 | S0－ss ${ }^{\text {c }}$ | 2 |
|  |  |  |  |  |  |  |  |  | 90－1L＊6 | ¢0－GZ•1 | G0－25．1 | S0－E8．1 | S0－02＇z | Z |
|  |  |  |  |  |  |  |  | 90－52．L | 90－5z•6 | G0－21．1 | ¢0－$\frac{1}{}$ ¢ 1 | ¢0－65 ${ }^{\text {¢ }}$ | S0－16．1 | 2 |
|  |  |  |  |  |  |  | 90－6t＇G | 90－56＊9 | 90－$\frac{1}{}$ ¢ 8 | 90－ $28 \cdot 6$ | 50－21： | ¢0－6を．1 | S0－29．L | 2 |
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|  |  |  |  |  | 90－¢て「غ | 90－90＊${ }^{\text {¢ }}$ | 90－18．7 | 90－29．5 | 90－95．9 | 90－89． 1 | 90－700．6 | 50－20．1 | S0－82： | 2 |
|  |  |  |  | 90－\＆ら． 2 | 90－91．${ }^{\text {c }}$ | $90-12 \cdot \varepsilon$ | 90－Lع＇t | 90－10．5 | 90－28．5 | 90－08．9 | 90－66： 2 | 90－2п－6 | S0－ 51.1 | 2 |
|  |  |  | 90－00＇Z | 90－Lt＇己 | 90－06＊$て$ |  | 90－98．$\varepsilon$ | 90－97．${ }^{\text {H }}$ | 90－81．5 | 90－70．9 | 90－01． 2 | 90－0t ${ }^{\text {c }} 8$ | 50－00．1 | 2 |
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|  | 90－12•1 | 90－95．1 | 90－28＊ | 90－80． 2 | 90－2E．2 | 90－1L 2 | 90－11．$\varepsilon$ | 90－85．$\varepsilon$ | 90－G1＇${ }^{\text {¢ }}$ | 90－78．＇． | 90－89．5 | 90－21．9 | 90－ $20 \cdot 8$ | \％ |
| 90－ $00 \cdot 1$ | 90－92＇t | 90－9t1 | 90－99＊ | 90－68．1 | 90－S1＇己 | 90－cti． | 90－18．2 | 90－$¢ 乙 \cdot \varepsilon$ | 90－til ${ }^{\text {c }}$ ¢ | 90－98．${ }^{\text {a }}$ | 90－11－G | 90－90．9 | 90－7C．${ }^{\text {90－8 }}$ | E |
| L0－E9．L | L0－25．8 | L0－ts： 6 | 90－20＇1 | 90－12．1 | 90－2E： 1 | 90－G5： | 90－8L． | 90－70． 2 | 90－9を． | 90－51． | 90－ 9 －${ }^{\text {a }}$ | ${ }^{90-\varepsilon 9} \cdot{ }^{\text {c }}$ | $90-81 \cdot \frac{1}{1}$ | \％ |
| L0－5¢．S | L0－56．5 | 10－79．9 | 20－Gt． | L0－8£．8 | 20－87\％ 6 | 90－80．1 | 90－ $10-\varepsilon \cdot 1$ | 90－1t．1 | 90－79． | 90－16： | 90－7\％${ }^{\text {9 }}$ | 90－26．1 | 90－9\％${ }^{\text {c }}$ | H |
| $\begin{aligned} & 20-86 \cdot \varepsilon \\ & 10-80 \cdot \varepsilon \end{aligned}$ |  | $20-76 \cdot$ <br> $0-28$. |  | L0－Eट．${ }_{\text {L }}$ | L0－tor |  | 20－90． | 20－11－8 |  | 90－60 | $90-62 \cdot 1$ | 90－25．1 | $90-\varepsilon 8^{\circ} \mathrm{L}$ | 5 |
| 62 | $88^{8}$ | L2 | 92 | ¢ | カट | £z | zz | 12 | 02 | 61 | 81 | L1 | 91 |  |
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|  |  |  |  |  |  |  |  |  |  | 20－Gtic | 20－2\％・サ | 20－10＊6 | 10－09．2 |  |
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|  |  |  |  |  |  |  |  | ع0－10＊\％ | 80－6п．9 | 20－70．1 | 20－18．1 | 20－59．E | 10－60： |  |
|  |  |  |  |  |  |  | 80－16．1 | ع0－66． 2 | ع0－95＇$\dagger$ | E0－52．L | 20－92：1 | 20－trs＇ 2 | 20－HE： |  |
|  |  |  |  |  |  | 10－06．6 | E0－0¢•1 | ع0－0C＇ट | ع0－0 $\varepsilon^{\prime} \varepsilon$ | ع0－mट．s | 80－60．6 | 20－78．1 | 20－ $2 \varepsilon^{.5}$ | 1 |
|  |  |  |  |  | to－stis 5 | 70－to ${ }^{\text {¢ }} 8$ | ع0－カt | ع0－¢9．1 | ع0－9力＇Z | ع0－16． | ع0－6L．9 | 20－8¢．1 | 20－00．${ }^{\text {\％}}$ | ， |
|  |  |  |  | 70－21．$\varepsilon$ | 70－15．h | to－ $2 ¢ .9$ | ＋0－28．8 | ع0－92＇1 | ع0－68．L | ع0－66＊ | ع0－12．5 | 20－90．1 |  | L |
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| S0－¢8．7 | 50－ 1.1 .9 | 50－ 52. | 50－18：6 | ＋0－92．1 | 70－59．1 | to－12． 2 | 70－70．$\underbrace{\text { ¢ }}$ | ＋0－7\％${ }^{\circ} \mathrm{H}$ |  | ع0－tio | ع0－78 |  | ع0－18．6 | 1 |
| 50－02．7 | 50－22．5 | S0－Es．9 | 50－52．8 | 170－90．1 | to－8E．L | 70－58．1 | 70－7¢ $+0-51 \cdot 己$ | $70-79$ <br> $70-80$ | to－9\％$\dagger 0-\varepsilon 9^{\circ}$ | H0－2n＊${ }^{8}$ | ¢0－$¢$ ¢ |  | ع0－85＇8 | 61 |
| S0－01．${ }^{\text {che }}$ | ¢0－18．${ }^{\text {c }}$ | ¢0－EL．${ }^{\text {c }}$ | 50－56． 5 | 50－19．${ }^{\text {c }}$ | 50－176．6 | ¢0－¢£． | ¢0－£8．1 | ＋0－ 0 ¢ $9 \cdot \frac{\Sigma}{}$ | $\quad 0-26 \cdot \varepsilon$ | ヶ0－En＇9 | ع0－91： |  |  | O2 |
| 50－89＇z | ¢0－82＇$\varepsilon$ | 50－90＊$\dagger$ | S0－01＇S | 50－85：9 | 50－25•8 | 70－カ1． | 70－85： | ＋0－92． z | п0－ 0 ¢ ${ }^{\text {¢ }}$ ¢ | to－85．5 | ع0－10．1 | ع0－81 ${ }^{\circ}$ | 80－89．9 | 12 |
| 50－2¢＇z | 50－¢8． 2 | S0－15＊ | S0－1ヵ＇H | S0－E9．5 | S0－9E． 2 | 50－98．6 | म0－9E． | п0－26＊ | 70－66． 2 | 70－68．7 | 70－16．8 | £0－76． | ع0－26．${ }^{5}$ | 寿 |
| S0－10．2 | S0－9力＇ | 50－50．${ }^{\text {c }}$ | S0－E8．$\varepsilon$ | G0－06．7 | 50－04．9 | 50－65 8 | 70－61： | 70－2L： | 70－29：2 |  | 70－16． | ع0－9\％． | ع0－18．${ }^{\text {c }}$ | \＆ |
| 50－92．1 | ¢0－st＇z | 50－99 ${ }^{\circ}$ | S0－п¢＇$\varepsilon$ | 50－82．7 | S0－09．5 | 50－85． 2 | カ0－50．1 | ＋0－25．！ | ＋0－28．2 | to－ $70-\varepsilon ¢$ | 70－98．9 | を0－で・1 |  | 52 |
| S0－S5． | 50－68．1 | 50－78． 2 | 50－76 ${ }^{\circ}$ | 50－LL | S0－76．${ }^{\text {c }}$ | 50－59．9 | 50－92．6 | ＋0－ヵ¢ + － | ＋0－20．2 | to－ $70-80 \cdot \frac{\varepsilon}{\varepsilon}$ | to－ 51.9 | ع0－62．t | ع0－50．\％ | 9 |
| 50－28． | 50－29： | 50－20．2 | 50－09．2 |  |  | 40－06：${ }^{\text {c }}$ | $50-72.8$ $50-2 E$. | ＋70－02 | ＋70－98．1 | ¢0－8L | 70－2Z• ${ }^{5}$ | ع0－Ll：1 | ع0－1L $\mathrm{\varepsilon}$ | 12 |
| 50－12． | $50-871$ | S0－78．1 | 40－18． | 50－26． | ¢0－06．$\frac{\varepsilon}{\varepsilon}$ | ¢0－EL－H | 50－ 59.9 | 50－02．6 | ＋10－05．l | ¢0－2G．${ }^{\text {¢ }}$ | т0－GL．$\dagger$ | £0－80．1 | ع0－1\％＇$\varepsilon$ | 82 |
| 50－80． 9 | ¢0－2¢ ${ }_{\text {c }}$ |  | 50－20．${ }^{\text {c }}$ | 50－68． | ¢0－65．${ }^{\text {c }}$ | 50－2L．t | G0－66． 5 | 50－8L．8 | to－2E： | to－0¢． 2 | H0－5¢ ${ }^{\text {¢ }}$ | 70－28．6 | ع0－G1．$¢$ | 62 |
| 90－GL ${ }^{9}$ | 50－20．1 | $50-\varepsilon{ }^{\text {c }}$－1 | 50－89．1 | S0－91． 2 | G0－58．2 | ¢0－18．$\varepsilon$ |  | ¢0－66． 2 | т0－пて 1 | 7\％－01＇ 2 | to－86．$\frac{1}{}$ | 70－80．6 | ع0－06： | OE |
| 90－55．5 | 90－08．9 | 90－2サ・8 | 50－20．1 | 50－8E．1 | ¢0－£8．1 | s0－05． 2 | S0－ES．$\varepsilon$ | 50－22． 5 | 50－61：8 | $70-6 \varepsilon$ | 70－29．2 | 70－91． | と0－66． | G8 |
| 90－G8．$\varepsilon$ | 90－EL＇t | 90－68．5 | 90－2t ${ }^{\circ}$ | 90－99．6 | 50－82： | 50－51． | 50－8t ${ }^{\circ}$ | $50-29 \cdot \%$ | 50－81．5 | 50－28．6 | ＋0－06． | 70－07． |  | 位 |
| 90－98＊2 | 90－15．$\varepsilon$ | 90－8¢ ${ }^{\circ}$ ¢ | 90－95．5 | 90－61．L | 90－75．6 | 50－08： | 50－58．1 | S0－12． | S0－2E．${ }^{\text {c }}$ | G0－07． | to－ $70-11.1$ | ¢0－1 $+0-8 \zeta$ |  |  |
| $\begin{gathered} 90-i z \cdot z \\ 51 \end{gathered}$ | $90-2 L \cdot 2$ | $\begin{gathered} 90-6 \varepsilon \cdot \varepsilon \\ \varepsilon! \end{gathered}$ |  | 90－25．5 | $90-6 \varepsilon^{\circ} \mathrm{L}$ 01 | S0－10．1 | SO－¢カ＇l | ${ }_{\text {S0－}}^{\text {L }}$ | ${ }_{9}^{50-9 \varepsilon}$ | $\mathrm{SO}_{5}$ |  |  |  |  |
| \＆t－0GL． | ＝ 9 d d | カレー | てサ「8＝5 | s2 |  |  | ら2－0カて | 1 ＝W3 |  | $-780 *$－ | 0y z－t | ع1－26 | ． $2=98$ |  |
|  |  |  |  |  |  |  | $0 L=0$ | 83 | ＊o | $70+00^{\circ} 1$ | 3N | $10+00 \cdot 1=$ |  | H |




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 can be readily carried out by assuming that the variation with temperature has a power-law form, $I\left(N_{\mathrm{e}}, T_{\mathrm{e}}, n^{\prime}, n\right)=A\left(N_{\mathrm{e}}, n^{\prime}, n\right) t^{\beta}$,where $t=10^{-4} T_{\mathrm{e}}(\mathrm{K})$ and $A\left(N_{\mathrm{e}}, n^{\prime}, n\right)$ and $\beta$ are constants. A similar procedure may be used to
interpolate in $n^{\prime}$ for $n^{\prime}>30$,
where $B\left(N_{\mathrm{e}}, T_{\mathrm{e}}, n\right)$ and $\gamma$ are constants. Variations with electron density are generally sufficiently
weak that linear interpolation suffices.

## 5.2 comparison with previous calculations

In the calculation of $b_{n}$ and $b_{n l}$ for hydrogenic ions, the rate coefficients for radiative processes are in principle known exactly, and in practice can be determined to any desired level of accuracy. We expect that these rates have an error of less than 1 per cent in our calculation. A similar degree of accuracy is obtained in numerical methods. The rates for collisional processes are far less well known, but the values of the $b$-coefficients are somewhat less sensitive to these processes. It is therefore of interest to compare our values of $b_{n}$ and $b_{n l}$ with those from other calculations with a different choice of collisional rate coefficients. We do, however, consider that the cross-sections of Percival \& Richards (1978), used in our calculations, are the best currently available. In the first stage calculation of $b_{n}$ we can compare our results with those of Brocklehurst (1970), who tabulates $b_{n}$ for H for $40 \leq n \leq 300$. We have made comparisons for $\log N_{\mathrm{e}}=3,4,5$ and $T_{\mathrm{e}}=5 \times 10^{3}$,


 density variation of 50 per cent a change in the collision rates of the same amount, which is


 that the differences between our results and those of Brocklehurst (1970) most probably arise from the different choice of cross-sections for the energy-changing collision rates
We can also make limited comparisons of our $b_{n}$ for H and those of Summers (1977). We
 , For $n \geq 40$, we find maximum differences of 2 per cent between our results and those of Summers. Summers' treatment of energy-changing collisions differs

 (Summers 1977) Summers used the impact parameter method (Seaton 1962; Burgess 1964), consistent with the cross-sections of Percival \& Richards (1978) used in this work. Again the
 although now the differences are in the opposite sense to those with respect to Brocklehurst

 larger differences are unexpected, in that the populations of such low-lying states depend
 disagreement for the $3 d$ and $4 d$ states is his rather sparse choice of $n$ states at low $n$ in the matrix condensation scheme (Summers 1986, personal communication)
Our results for the relative intensities and effective recombination coefficients can be

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 compared with those of Brocklehurst (1971) for H and $\mathrm{He}^{+}$, and with those of Seaton (1978) for








 decreasing with rising electron temperature.
We find slightly larger differences when comparing with the work of Seaton (1978), who obtained relative line intensities for $\mathrm{He}^{+} n \rightarrow 2$ and $n \rightarrow 3$ series. These intensities were derived from the zero-density calculations of Pengelly (1964) and the density-dependent calculations of Brocklehurst (1971) by using scaling laws. We differ from Seaton by a maximum of 6 per cent, with an average difference of 3 per cent.
Brocklehurst (1971) also tabulates values for the effective recombination coefficients for the H $4 \rightarrow 2$ and $\mathrm{He}^{+} 4 \rightarrow 3$ transitions at a few temperature and densities. We reproduce his results for
these quantities to within 1 per cent for all cases.

## Acknowledgments

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818
 Appendix: Rate coefficients for electron induced transitions in $\mathrm{H}_{\text {and }} \mathrm{He}^{+}$ Consider two atomic states $i$ and $j$. The collision strength $\Omega(i, j)$ for excitation from $i$ to $j$ is related
(A1)
where $\omega_{i}$ is the statistical weight of the initial state, $E_{i}$ is the energy of the incident electron relative
to the initial state and $a_{0}$ is the Bohr radius. The rate coefficient for excitation from $i$ to $j, q_{i j}$, may to the initial state and $a_{0}$ is the Bohr radius. The rate coefficient for excitation from $i$ to $j, q_{i j}$, may
be written

[^2]where $\Delta E_{i j}$ is the energy of the transition and $Y\left(i, j, T_{\mathrm{e}}\right)$ is an effective collision strength, obtained by averaging $\Omega(i, j)$ over a Maxwellian distribution of electron velocities,

## (A3)

We require values of $Y$ for transitions between the $n=1,2$ and 3 states of H and $\mathrm{He}^{+}$. We give the sources of this information below.
We use the effective collision strengths of Aggarwal (1983), who makes the fit
$Y\left(T_{\mathrm{e}}\right)=A+B T_{\mathrm{e}}+C T_{\mathrm{e}}^{2}+D T_{\mathrm{e}}^{3}$
and tabulates $A, B, C, D$ for each transition. For $1 s \rightarrow 2 s$ and $1 s \rightarrow 2 p$, Aggarwal uses the collision strengths recommended by Callaway \& McDowell (1983). For the $1 \rightarrow 3$ transitions he uses the




[^3]Recombination-line intensities for hydrogenic ions 819
about 0.04 Ryd of the $n=2$ threshold, there are considerable uncertainties in the rates derived, treated in the unitarized distorted-wave method. The use of the distorted-wave method for $L>12$
 Fit from Callaway (1983).
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\underset{\sim}{n} & \underset{\sim}{n} & 0 \\
\dot{0} & \dot{0} & \dot{0}
\end{array}
$$
\] particularly at higher temperatures.

We have carried out a six-state calculation ( $1 s, 2 s, 2 p, 3 s, 3 p, 3 d$ ) of collision strengths for $\mathrm{He}^{+}$in the energy region above the $n=3$ threshold. We define $E_{1}$ to be the energy (in Ryd) of the incident electron relative to the ground state of $\mathrm{He}^{+}$, and $L$ to be the total orbital angular momentum of the $\mathrm{He}^{+}+e^{-}$system. We calculate collision strengths at the energies $E_{1}=3.57,3.8$, $4.0,4.41,5.0$ and 6.0 Ryd , for $L \leq 23$. For $0 \leq L \leq 12$ we have used the close-coupling code impact (Crees, Seaton \& Wilson 1978), and for $13 \leq L \leq 23$, the unitarized distorted-wave method described by Eissner \& Seaton (1972). Comparison of collision strengths calculated using the two methods, at $E_{1}=6$ Ryd and $L=12$ shows that the agreement is good with the exception of the $2 s \rightarrow 3 p$ and $2 s \rightarrow 3 d$ transitions, where the difference is 13 per cent in both cases. This difference is due to strong coupling between the channels based on the $n=3$ states, which is not correctly implies a maximum uncertainty from this source of 4 per cent for these two transitions. The error




 lower energies, such corrections are negligible ( $<1$ per cent).
Our calculated inelastic scattering collision strengths are given in Table A1. More elaborate
 $1 s \rightarrow 2 p$ collision strengths at intermediate energies. We adopt the following values for these
transitions
$1 s \rightarrow 2 s \quad 3.0 \leq E_{1} \leq 3.24 \quad$ Linear interpolation of calculation of Morgan (1980).
Linear interpolation of calculation of Wakid \& Callaway (1980). $3.24<E_{1} \leq 3.556$
$3.556<E_{1} \leq 4.0$
 $4.5<E_{1} \leq \infty$
Table A1. Collision strengths, $\Omega$, for $\mathrm{He}^{+}+e^{-}$

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D. G. Hummer and P.J. Storey
$1 s \rightarrow 2 p \quad 3.0 \leq E_{1} \leq 3.24 \quad$ Linear interpolation of calculation of Morgan (1980). $\begin{array}{ll}3.24<E_{1} \leq 3.556 & \Omega=0.279+0.357\left(E_{1}-3\right) . \\ 3.556<E_{1} \leq 4.0 & \Omega=0.262+0.311\left(E_{1}-3\right) .\end{array}$ $4.0<E_{1} \leq \infty \quad$ Fit from Callaway (1983).
The calculation of Morgan (1980) is the most elaborate in the near-threshold region. In the energy region $3.24<E_{1} \leq 3.556$ we use the results of the calculation by Burke \& Taylor (1969) who use a three-state close coupling expansion with correlation terms. These results contain considerable resonance structure. The value of $\Omega(1 s, 2 s)=0.147$ in this energy region is a simple mean of their results, whereas the linear relation for $\Omega(1 s, 2 p)$ is a least-squares fit to their results, this form being chosen to reflect the shape of the experimental curve of Daschenko et al. (1974) in this energy region. Between the $n=3$ threshold and the ionization limit, $3.556<E_{1} \leq 4.0$, we are guided by our own six-state close-coupling calculation. We consider that this calculation gives the most reliable result for the collision strengths immediately above the $n=3$ threshold, but that it increases too rapidly with increasing energy due to unaccounted for open channels. The choice of
 (1980) at the ionization limit, whereas the linear behaviour of $\Omega(1 s, 2 p)$ is chosen to connect the six-state close-coupling result at the $n=3$ threshold with the Callaway (1983) fit at the ionization
 collision strengths $\Upsilon\left(T_{\mathrm{e}}\right)$ have been calculated (equation A3) and fitted to the form,

VOL. 224, NO. 4, 1987
Recombination-line intensities for hydrogenic ions - I.
Case $B$ calculations for $H I$ and HeII

## D.G. Hummer and P.J. Storey



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$\therefore \therefore \therefore=\therefore$－Nimininnion




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Nm＝NーMーNNNMmin


































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MツツN゙がかっか


－MomN－Mmo



























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둥․



























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＝MMMNNNーーロ


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    - \(-m-n=-\dot{n}\)



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天系商会
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シニュのーか

ジージー
シューががッ

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Mのmがべか


シーツッががm人min －－mini－s
戸mmminnim

－\(-n=-\) Nim－a


\(=ゅ-m \infty-n=-\)
シーツッツツッツがの













crnmes～ors－om－9
が戸ジョジミジ


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がががシニシニジ



 －Nmzinn

 NNminaー－－－－N NNmmininio











 －Nos



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- Nim=
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人



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人

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がごか
の上よ



のッジャッ
天
フー～ール゙


T 二isions


守べローシ
ショシッスがッロ

\(\therefore \mathrm{G}=\mathrm{M}=\mathrm{m}-\mathrm{mN}\)












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Эロニニベかめx



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MOMMNNNNNN






 ¢シーデMッチーーか




















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ショシーン

mz゙心ース

\(9 m+0 \infty 0 \rightarrow\)

Zロ0～～

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二ロニデー
べッジッヂか










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スペロN







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－©ox





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ジミツッチ

シューーツ
ミニニがが

シーローがら
シニジッヂか

mivivimi
シニショーロが
 mm＝nヶcm ジジッジッヂが

NAMmがMiN
シニジロジッツッチ
＂
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人


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웅엇웅


\(\stackrel{n}{5}\)





 ジジッーロ Faind がき心． yininon－





\(\dot{O}-\operatorname{mininc}\)
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 －－inハースース＝－




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も三品



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人
N～ジシーシ

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N
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m＝\(\dot{6}\)－
人
三nin－…ninin

－bérinniñN

－ig－inivimimi

\section*{}


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＝



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minnop

守守守
－mのベが
天象象家

MMMシーローか

デテテーディ
戸ッツツNがー

品品分

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人
－ \(0=\)－Nvinivim


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- \(\because \cdots\) -


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z z onnmmin=ino onn
象象

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\(\therefore \therefore=-2=\)
ニキミ゚


＝－m～N
\(=0\)－incincm
N

\(\therefore \therefore \operatorname{Anmmin}=\sin\)
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\(\stackrel{5}{6}\)

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ショニッNN～ーロ

～ゥoncisoñ

















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\(\stackrel{\sim}{\sim} \dot{\sim}\)
 NM＝

人家定号


miso

ジローシー・ー～
气 mini－i－ninin





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 \(\therefore\)－－Mrimin＝indin
－ ＝nmozono


－
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\section*{ジッデー}

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도ำํํ
ローニーへm








がニシmmminm
－
ற்ーシーシーシーか


シーティローズージー
がニニジッツッツが

\(\therefore \therefore \therefore \dot{A}-\mathrm{Nanc}\)








\(-\therefore-\sin -\therefore=-0\)



－ジミニジニ






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๓足
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- Nonmmin




A思


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ゅ品
ㅋํ웁우웅

\(\therefore \therefore\)－


 z \(-\infty\)
今心のNう












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ジッグがか

守曻
戸戸ががか

テーデッ
ショウッベが



 テテテビッフ





○べ八テティ～


























－armonson
ま







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人
－Nimis



NOOMO

도NNNN






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\(0-0 n \infty=0\) ．

amnoraco



促



\(\therefore\)－mininuo


ジッペか







































\section*{}

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z \(\stackrel{\rightharpoonup}{6}\) \(\stackrel{\triangleleft}{¿}\)

为
－nmin
为会会 \(\therefore \sim \sim \circ 0^{\circ}\)


为 か－Nin＝
















O定























\(\stackrel{\rightharpoonup}{\vec{x}}\)

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のin 2
BNomio
がぁがめ゙か ヘ上こ゚心2 ジデッロ ぶッチッN゙

デテーがロ







天



テーデッチーディ



























 NNoNmiran








人


 －＝No

 OSOM




 ～品 き

\section*{ \\ －nimini}

R
\(\sin =0 \infty^{\infty}\)

जhín wig

NMini＝
N －－iomz
 mioniono
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 －
 응융 ๙ 2人）































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\(\dot{m}=0-\cdots N\) NN

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N


がm゙ローシー
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ごッ゙ー・ーシ
a
- ジーシージ
MiN N-NNNN

-


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 ＝O～NM，FNOR



 －Moms nionco

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 ．\(\because\) •





ジミニ゙

ヴーシーシ


\(\therefore=-=-\)
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C
nmor：＝ーの

ニン日ぎ心
－morーテーの
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に高边
－ッニーかーツース










so－ri－arimioire













 －\(\therefore=-\mathrm{Namisio}\)






－ 0



＊ －iw－ninmoniono nirmina









 Mース象象
象 \(\therefore\)－人 －00－ncommi＝
 ヘNジ心ッロッシース







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\(\therefore\) añosion



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のミミース
呺空
\(\therefore \div \div\)
ニシニシミミ

\(\therefore \therefore \div \div\)
コーニ゙ラシミス

－\(-7 \div\)
づニニーシニミ

－〒－テ二



シえミワのニニシミス


シミニシニロッニニミか







＝\(\because=-\)－














－moーniviza＝


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禺会気
人
\(\therefore\) 电
为
为象
\(=-\equiv 0\)


ニ上権



－－－－•••・ーシ
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必电













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ジジが気

－－スィーテ
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このジッジが


こここのらえ

のニーがいいい
ニニニニミごこニミ


ニこニニニミこえここ


ここここここご二こここミ
ここごここここここご

ニッニニニこニこミこここ
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ミニニニニニンーごこここう

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= 人












    尼总







MAN天NNo


气气 じ1

 ニ







\(\therefore\) an 2 ninusion

 NOONO




읍气ロロロ
ละ，
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象领 \(\therefore \therefore \dot{\sim}\)气ききききご会



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옹․








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戸ージ品

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－ヲテッチ
シニーニッス゚


がニッシ゚ロッ

シーッダラジ

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へのニーシミニニッ゚



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の会会领気気
ニ5

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Qx




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ベッロース
天上定高
ベヒーシー


\(\therefore=\therefore\) ○ー


－mッーデテ


－sincio？
シーががニシべ


のックローニシシャッ






 A－





















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＝i－coin






 Smaño mo－＝min－

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デロ


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－ームデーーNNmのローmー・











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or


のジロニッ
－－＝～m
シーテのが心
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へーmーニ～～

べずべN
nanamen
がバ戸戸がNべN
 シーーローーツN

 mio－misi－min
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Nへーシーロ0090




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－－mi＝a
Nがmaが？

noinimomo


sr－nn－ma

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Nin－N－Ni＝cinN
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～ N二Nーツ－nmon

Non No
Nimo－niciomini－





－Nm由Nm＝0の－NEー－


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FMmmmNN№

ローテーテデッロッチ

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[^0]:    

[^1]:    (zz)
    where $N_{21}$ depends parametrically on $N_{\mathrm{c}}^{\mathrm{c}}$. In general equation (22) is quadratic in $N_{\mathrm{e}}^{\mathrm{c}}$, although it
     regions (2) and (3), for which the condition $3 N_{e} C \leqslant A_{2 q}$ is valid. if, in addition

[^2]:    $\Upsilon\left(i, j, T_{\mathrm{e}}\right) \exp \left(-\Delta E_{i j} / k T_{\mathrm{e}}\right) \mathrm{cm}^{3} \mathrm{~s}^{-1}$

[^3]:    

[^4]:    
    

[^5]:    $\stackrel{s}{3}$

