Second-order gravitational self-force

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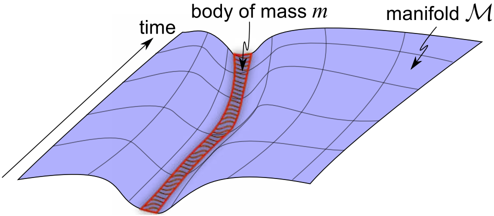
A small extended body moving through spacetime

Fundamental question

- how does a body's gravitational field affect its own motion?

Regime: asymptotically small body

- examine spacetime $(\mathcal{M}, g_{\mu\nu})$ containing body of mass m and external lengthscales $\mathcal{R}$
- seek representation of motion in limit $\epsilon = m/\mathcal{R} \ll 1$

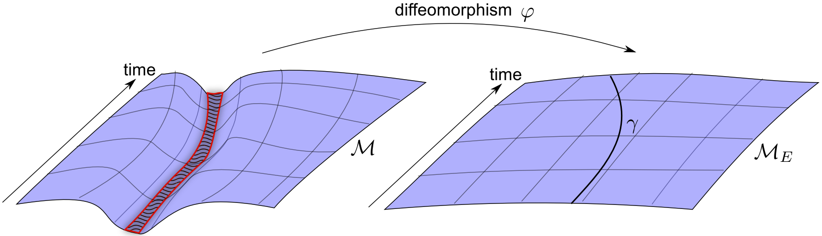


Gravitational self-force

- treat body as source of perturbation of external background spacetime $(\mathcal{M}_E, g_{\mu\nu})$

$$\mathsf{g}_{\mu\nu} = g_{\mu\nu} + \epsilon h^{(1)}_{\mu\nu} + \epsilon^2 h^{(2)}_{\mu\nu} + \dots$$

- $h^{(n)}_{\mu\nu}$ exerts *self-force* on body
- self-force at linear order in ϵ first calculated in 1996 [Mino, Sasaki, and Tanaka], now on firm basis [Gralla & Wald; Pound; Harte]

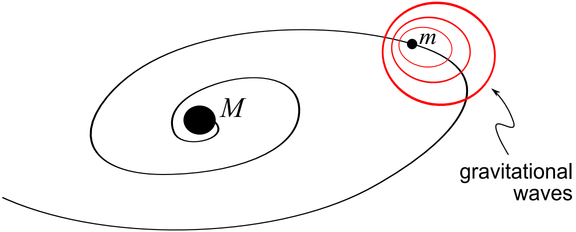


Body in exact spacetime

Representation of motion in external spacetime

Canonical example: extreme-mass-ratio inspiral

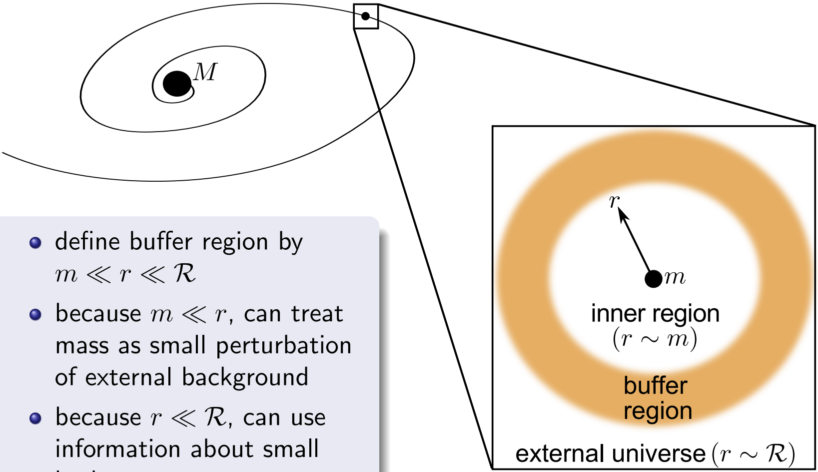
- solar-mass neutron star or black hole orbits supermassive black hole
- m = mass of smaller body, $\mathcal{R} \sim M$ = mass of large black hole
- $(\mathcal{M}_E, g_{\mu\nu})$ = Kerr spacetime of large black hole



Why second order?

- inspiral occurs very slowly, on timescale $1/\epsilon$
  $\Rightarrow$ need $O(\epsilon^2)$ terms in acceleration to get trajectory correct at $O(1)$
- also useful to complement PN and NR

How to determine motion: buffer region

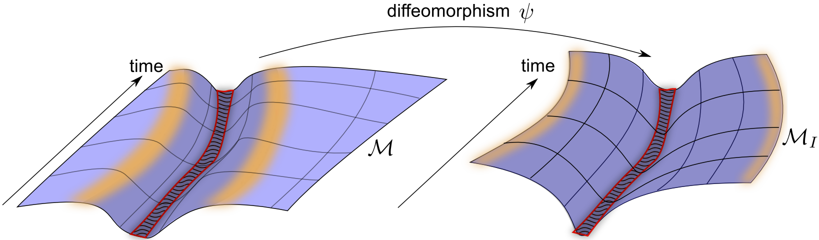


- define buffer region by $m \ll r \ll \mathcal{R}$
- because $m \ll r$, can treat mass as small perturbation of external background
- because $r \ll \mathcal{R}$, can use information about small body

Matched asymptotic expansions: *inner expansion*

Zoom in on body

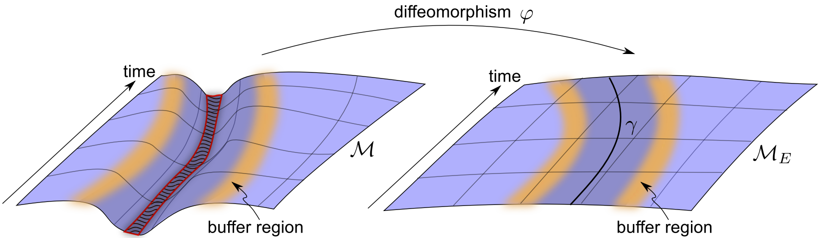
- map ψ keeps size of body fixed, sends other distances to infinity (e.g., using coords $\sim r/\epsilon$)
- unperturbed body defines background spacetime $g_{I\mu\nu}$ in inner expansion
- buffer region at asymptotic infinity
  $\Rightarrow$ can define multipole moments



Matched asymptotic expansions: *outer expansion*

Send body to zero size around a worldline

- map φ shrinks body to zero size, holding other distances fixed
- build metric $g_{\mu\nu} + \epsilon h^{(1)}_{\mu\nu} + \epsilon^2 h^{(2)}_{\mu\nu} + \ldots$ in external universe (outside buffer region) subject to *matching condition*: in coords centered on γ, metric in buffer region must agree with inner expansion



Metric in buffer region

Expansion for small r

- presence of *any* compact body in inner region leads to

$$h^{(1)}_{\mu\nu} = \frac{1}{r}h^{(1,-1)}_{\mu\nu} + h^{(1,0)}_{\mu\nu} + rh^{(1,1)}_{\mu\nu} + O(r^2)$$

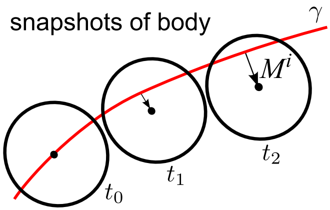
$$h^{(2)}_{\mu\nu} = \frac{1}{r^2}h^{(2,-2)}_{\mu\nu} + \frac{1}{r}h^{(2,-1)}_{\mu\nu} + h^{(2,0)}_{\mu\nu} + O(r)$$

  where r is distance from γ

- most divergent terms are background spacetime in inner expansion: $g_{I\mu\nu} = \eta_{\mu\nu} + \frac{1}{r}h^{(1,-1)}_{\mu\nu} + \frac{1}{r^2}h^{(2,-2)}_{\mu\nu} + O(1/r^3)$

Relating worldline to body

- define γ to be worldline of body iff mass dipole terms vanish in coords centered on γ



Solving the EFE with an accelerated source

Expansion of EFE

- allow γ to depend on ϵ and assume outer expansion of form

$$\mathsf{g}_{\mu\nu}(x,\epsilon) = g_{\mu\nu}(x) + h_{\mu\nu}(x;\gamma)$$
$$= g_{\mu\nu}(x) + \epsilon h^{(1)}_{\mu\nu}(x;\gamma\ ) + \epsilon^2 h^{(2)}_{\mu\nu}(x;\gamma\ ) + \ldots$$

- need a method of systematically solving for each $h^{(n)}_{\mu\nu}$
  $\Rightarrow$ impose Lorenz gauge on total perturbation: $\nabla_\mu \bar{h}^{\mu\nu} = 0$
- linearized Einstein tensor $\delta G_{\mu\nu}$ becomes a wave operator and EFE becomes a weakly nonlinear wave equation:

$$\Box \bar{h}_{\mu\nu}[\gamma] + 2R_\mu{}^\rho{}_\nu{}^\sigma \bar{h}_{\rho\sigma}[\gamma] = 2\delta^2 G_{\mu\nu}[h] + \ldots$$

  (no stress-energy tensor because equation written outside body)
- can be split into wave equations for each subsequent $h^{(n)}_{\mu\nu}[\gamma]$ and exactly solved for arbitrary γ
- $\nabla_\mu \bar{h}^{\mu\nu} = 0$ determines acceleration of γ

Solving the EFE with an accelerated source

Expansion of EFE

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$$\begin{aligned} \mathrm{g}_{\mu\nu}(x,\epsilon) &= g_{\mu\nu}(x) + h_{\mu\nu}(x;\gamma) \\ &= g_{\mu\nu}(x) + \epsilon h^{(1)}_{\mu\nu}(x;\gamma_\epsilon) + \epsilon^2 h^{(2)}_{\mu\nu}(x;\gamma_\epsilon) + \ldots \end{aligned}$$

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General solution in buffer region

First order

- field naturally splits in two: $h^{(1)}_{\mu\nu} = h^{S(1)}_{\mu\nu} + h^{R(1)}_{\mu\nu}$
- $h^{S(1)}_{\mu\nu} \sim 1/r + \ldots$ defined by mass monopole m
- $h^{R(1)}_{\mu\nu} \sim r^0 + \ldots$ undetermined homogenous solution regular at $r = 0$
- $\nabla_\mu \bar{h}^{\mu\nu} = 0 \Rightarrow \dot{m} = 0$ and $a^\mu_{(0)} = 0$

Second order

- field naturally splits in two: $h^{(2)}_{\mu\nu} = h^{S(2)}_{\mu\nu} + h^{R(2)}_{\mu\nu}$
- $h^{S(2)}_{\mu\nu} \sim 1/r^2 + 1/r + \ldots$ defined by
  1. mass correction δm
  2. mass dipole M^μ (set to zero with appropriate choice of γ)
  3. spin dipole S^μ
- $\nabla_\mu \bar{h}^{\mu\nu} = 0 \Rightarrow \dot{S}^\mu = 0$, $\delta\dot{m} = \ldots$, and $a^\mu_{(1)} = \ldots$

Matching to an inner expansion

Inner expansion

- could continue with same method to find $a^{\mu}_{(2)}$ from $h^{(3)}_{\mu\nu}$
- instead, get more information from inner expansion
- assume metric in inner expansion is Schwarzschild as tidally perturbed by external universe
- write tidally perturbed Schwarzschild metric in mass-centered coordinates

Matching

- expand inner metric in buffer region (i.e., for $r \gg m$)
- demand inner and outer expansions in buffer region are related by unique gauge transformation $x^{\mu} \to x^{\mu} + \epsilon\xi^{\mu} + \ldots$
- restrict gauge transformation to include no translations at $r = 0$ to ensure worldline correctly associated with center of mass

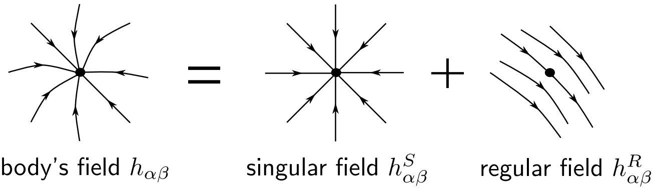
Equation of motion

Self-force

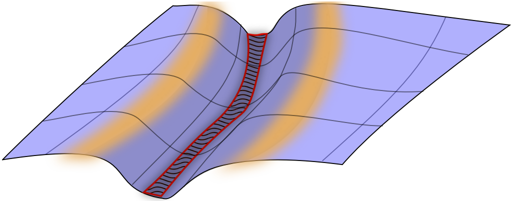
- matching procedure yields acceleration

$$a^{\mu} = \frac{1}{2}\left(g^{\mu\nu} + u^{\mu}u^{\nu}\right)\left(g_{\nu}{}^{\rho} - h^{\mathrm{R}\rho}_{\nu}\right)\left(h^{\mathrm{R}}_{\sigma\lambda;\rho} - 2h^{\mathrm{R}}_{\rho\sigma;\lambda}\right)u^{\sigma}u^{\lambda} + O(\epsilon^3)$$

  where $a^{\mu} = a^{\mu}_{(0)} + \epsilon a^{\mu}_{(1)} + \epsilon^2 a^{\mu}_{(2)} + \ldots$
  and $h^{R}_{\mu\nu} = \epsilon h^{R(1)}_{\mu\nu} + \epsilon^2 h^{R(2)}_{\mu\nu} + \ldots$
- this is geodesic equation in metric $g_{\mu\nu} + h^{R}_{\mu\nu}$
- equation for more generic body will be the same, modified only by body's multipole moments



Obtaining global solution



Puncture/effective-source scheme

- define $h^{\mathcal{P}}_{\mu\nu}$ as small-r expansion of $h^{\mathrm{S}}_{\mu\nu}$ truncated at order r or higher
- define $h^{\mathcal{R}}_{\mu\nu} = h_{\mu\nu} - h^{\mathcal{P}}_{\mu\nu} \simeq h^{\mathrm{R}}_{\mu\nu}$

The point...

- $h^{\mathrm{S}}_{\mu\nu}$ found in buffer region suffices to determine both $h^{\mathrm{R}}_{\mu\nu}$ and global solution outside body

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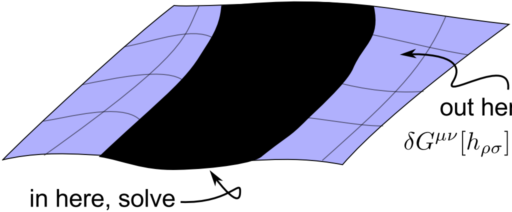

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Summary

Determining the motion of a small body

- define a worldline of an asymptotically small body, even a black hole, by comparing metric in a buffer region around body in full spacetime and in background spacetime
- determine equation of motion from consistency of Einstein's equation

Future work

- find equation for spinning, non-spherical body
- implement puncture scheme