

3-22-2016

SN 2012cg: Evidence for Interaction Between a Normal SN Ia and a Non-degenerate Binary Companion

G. H. Marion

Peter J. Brown

Jozsef Vinkó

Jeffrey M. Silverman

David J. Sand

See next page for additional authors

Follow this and additional works at: https://cedar.wvu.edu/physicsastronomy_facpubs

Part of the [Astrophysics and Astronomy Commons](#)

Recommended Citation

Marion, G. H.; Brown, Peter J.; Vinkó, Jozsef; Silverman, Jeffrey M.; Sand, David J.; Challis, Peter; Kirshner, Robert P.; Wheeler, J. Craig; Berlind, Perry; Brown, Warren R.; Calkins, Michael L.; Camacho, Yssavo; Dhungana, Govinda; Foley, Ryan J.; Friedman, Andrew S.; Graham, Melissa L.; Howell, D. Andrew; Hsiao, Eric Y.; Irwin, Jonathan M.; Jha, Saurabh W.; Kehoe, Robert; Macri, Lucas M.; Maeda, Keiichi; Mandel, Kaisey; McCully, Curtis; Pandya, Viraj; Rines, Kenneth J.; Wilhelmy, Steven; and Zheng, Weikang, "SN 2012cg: Evidence for Interaction Between a Normal SN Ia and a Non-degenerate Binary Companion" (2016). *Physics & Astronomy*. 24.

https://cedar.wvu.edu/physicsastronomy_facpubs/24

Authors

G. H. Marion, Peter J. Brown, Jozsef Vinkó, Jeffrey M. Silverman, David J. Sand, Peter Challis, Robert P. Kirshner, J. Craig Wheeler, Perry Berlind, Warren R. Brown, Michael L. Calkins, Yssavo Camacho, Govinda Dhungana, Ryan J. Foley, Andrew S. Friedman, Melissa L. Graham, D. Andrew Howell, Eric Y. Hsiao, Jonathan M. Irwin, Saurabh W. Jha, Robert Kehoe, Lucas M. Macri, Keiichi Maeda, Kaisey Mandel, Curtis McCully, Viraj Pandya, Kenneth J. Rines, Steven Wilhelmy, and Weikang Zheng



SN 2012cg: EVIDENCE FOR INTERACTION BETWEEN A NORMAL SN Ia AND A NON-DEGENERATE BINARY COMPANION

G. H. MARION^{1,2}, PETER J. BROWN³, JOZSEF VINKÓ^{1,4}, JEFFREY M. SILVERMAN^{1,21}, DAVID J. SAND⁵, PETER CHALLIS², ROBERT P. KIRSHNER², J. CRAIG WHEELER¹, PERRY BERLIND², WARREN R. BROWN², MICHAEL L. CALKINS², YSSAVO CAMACHO^{6,7}, GOVINDA DHUNGANA⁸, RYAN J. FOLEY^{9,10}, ANDREW S. FRIEDMAN^{2,11}, MELISSA L. GRAHAM¹², D. ANDREW HOWELL^{13,14}, ERIC Y. HSIAO^{15,16}, JONATHAN M. IRWIN², SAURABH W. JHA⁶, ROBERT KEHOE⁸, LUCAS M. MACRI³, KEIICHI MAEDA^{17,20}, KAISEY MANDEL², CURTIS MCCULLY¹³, VIRAJ PANDYA^{6,18}, KENNETH J. RINES¹⁹, STEVEN WILHELMI¹⁹, AND WEIKANG ZHENG¹²

¹ University of Texas at Austin, 1 University Station C1400, Austin, TX, 78712-0259, USA

² Harvard-Smithsonian Center for Astrophysics, 60 Garden St., Cambridge, MA 02138, USA; ghmarion@gmail.com

³ George P. and Cynthia Woods Mitchell Institute for Fundamental Physics & Astronomy, Texas A. & M. University, Department of Physics and Astronomy, 4242 TAMU, College Station, TX 77843, USA

⁴ Department of Optics and Quantum Electronics, University of Szeged, Domter 9, 6720, Szeged, Hungary

⁵ Physics Department, Texas Tech University, Lubbock, TX 79409, USA

⁶ Department of Physics and Astronomy, Rutgers the State University of New Jersey, 136 Frelinghuysen Road, Piscataway, NJ 08854 USA

⁷ Department of Physics, Lehigh University, 16 Memorial Drive East, Bethlehem, Pennsylvania, PA 18015, USA

⁸ Department of Physics, Southern Methodist University, Dallas, TX 75275, USA

⁹ Astronomy Department, University of Illinois at Urbana-Champaign, 1002 W. Green Street, Urbana, IL 61801 USA

¹⁰ Department of Physics, University of Illinois Urbana-Champaign, 1110 W. Green Street, Urbana, IL 61801 USA

¹¹ Center for Theoretical Physics and Department of Physics, Massachusetts Institute of Technology, Cambridge, MA 02139, USA

¹² Department of Astronomy, University of California, Berkeley, CA 94720-3411, USA

¹³ Las Cumbres Observatory Global Telescope Network, 6740 Cortona Drive, Suite 102, Goleta, CA 93117, USA

¹⁴ Department of Physics, University of California, Santa Barbara, Broida Hall, Mail Code 9530, Santa Barbara, CA 93106, USA

¹⁵ Carnegie Observatories, Las Campanas Observatory, Colina El Pino, Casilla 601, Chile

¹⁶ Department of Physics and Astronomy, Aarhus University, Ny Munkegade 120, DK-8000 Aarhus C, Denmark

¹⁷ Department of Astronomy, Kyoto University, Kitashirakawa-Oiwake-cho, Sakyo-ku, Kyoto 606-8502, Japan

¹⁸ Department of Astrophysical Sciences, Peyton Hall, Princeton University, Princeton, NJ 08544, USA

¹⁹ Department of Physics & Astronomy, Western Washington University, 516 High Street, Bellingham, WA 98225, USA

²⁰ Kavli Institute for the Physics and Mathematics of the universe (WPI), University of Tokyo, 5-1-5 Kashiwanoha, Kashiwa, Chiba 277-8583, Japan

Received 2015 July 26; accepted 2016 February 8; published 2016 March 22

ABSTRACT

We report evidence for excess blue light from the Type Ia supernova (SN Ia) SN 2012cg at 15 and 16 days before maximum B -band brightness. The emission is consistent with predictions for the impact of the supernova on a non-degenerate binary companion. This is the first evidence for emission from a companion to a normal SN Ia. Sixteen days before maximum light, the $B - V$ color of SN 2012cg is 0.2 mag bluer than for other normal SN Ia. At later times, this supernova has a typical SN Ia light curve, with extinction-corrected $M_B = -19.62 \pm 0.02$ mag and $\Delta m_{15}(B) = 0.86 \pm 0.02$. Our data set is extensive, with photometry in seven filters from five independent sources. Early spectra also show the effects of blue light, and high-velocity features are observed at early times. Near maximum, the spectra are normal with a silicon velocity $v_{Si} = -10,500 \text{ km s}^{-1}$. Comparing the early data with models by Kasen favors a main-sequence companion of about six solar masses. It is possible that many other SN Ia have main-sequence companions that have eluded detection because the emission from the impact is fleeting and faint.

Key words: supernovae: general – supernovae: individual (2012cg)

1. INTRODUCTION

It is widely accepted that Type Ia supernovae (SN Ia) are the thermonuclear explosions of carbon–oxygen white dwarfs, and many of them appear to explode near the Chandrasekhar mass (M_{Ch} ; e.g., Hillebrandt & Niemeyer 2000), though they may arise from progenitors of other masses as well (e.g., Scalzo et al. 2014). Two general progenitor scenarios are commonly considered for a white dwarf to accrete sufficient mass to approach the Chandrasekhar limit. In the single degenerate (SD) model, a non-degenerate binary companion star deposits matter onto a white dwarf. As the white dwarf nears the Chandrasekhar mass, a thermonuclear runaway is initiated (Whelan & Iben 1973; Nomoto 1982). The double degenerate (DD) scenario postulates that two carbon–oxygen white dwarfs will merge via gravitational inspiral and explode by subsequent

carbon ignition (Iben & Tutukov 1984; Webbink 1984). Other models are being explored, such as one in which the triggering mechanism of the SN Ia explosion is the head-on collision of two white dwarfs in a three-body system (Kushnir et al. 2013) and the core degenerate scenario that involves the merger of a white dwarf and the hot core of a massive asymptotic giant branch star (Kashi & Soker 2011).

There is some evidence that both the SD and DD scenarios contribute to the SN Ia population (see Maoz et al. 2014, for a recent review). For example, *Hubble Space Telescope* (HST) deep pre-explosion imaging of the site of SN 2011fe rules out evolved companion stars with $M > 3.5M_{\odot}$ (Li et al. 2011; Graur et al. 2014). This result does not rule out the SD scenario, even for this individual case, but it does cut a swath through the allowable parameter space. On the other hand, the SN Ia PTF11kx had clear signs of interaction with shells of circumstellar medium (CSM). The details suggest a SD system

²¹ NSF Astronomy and Astrophysics Postdoctoral Fellow.

with a Red Giant companion star in a symbiotic nova configuration (Dilday et al. 2012).

Many studies have searched for clues indicating the interaction between a normal SN Ia and CSM, presumably pointing to the SD scenario, but they are rarely as decisive as PTF11kx. Detection of interaction based on variable Na I features (e.g., Patat et al. 2007; Blondin et al. 2009; Sternberg et al. 2011; Foley et al. 2012b; Maguire et al. 2013), and measurements of high velocity features (HVF) from multiple spectral lines (e.g., Marion et al. 2013; Silverman et al. 2015) have all been used to investigate potential SD systems. See Maoz et al. (2014) for a more thorough discussion of work in this area.

A clearer signpost of the SD scenario lies in the very early light curves (LC). A nearby non-degenerate binary companion will encounter the shock wave and the expanding debris from the explosion. The interaction compresses and heats matter at the point of impact, but not all of the thermal energy is emitted in a prompt burst. Deeper layers of the ejecta continue to impact the companion and raise the local temperature. The extra luminosity from this interaction will be strongest in the ultraviolet and blue optical bands, and it will only be detectable for a few days after explosion (Kasen 2010, but see Maeda et al. 2014; Kutsuna & Shigeyama (2015)).

The effect on the observed brightness depends on the viewing angle (Kasen 2010; Brown et al. 2012). Interaction is unlikely to be detected if the impact location is too far from the direct line of sight to the SN. Thus, not all explosions in SD systems are expected to produce an early LC signal; perhaps only $\sim 10\%$ of cases will do so (Kasen 2010).

No previous detections have been reported for the interaction between a *normal* SN Ia and its companion after inspections of hundreds of LC from SN Ia, though many of these were not obtained early enough to test the Kasen (2010) model predictions.

Hayden et al. (2010) looked for interaction signals in *B*-band LC of 108 SDSS SN Ia. They found that companion stars would have to be less than about $6M_{\odot}$ on the main-sequence and they strongly disfavored Red Giant companions. Ganeshalingam et al. (2011) also examined *B*-band rise time behavior for 61 SN Ia and found no evidence of companion interaction. Tucker (2011) analyzed *U*-band LCs of ≈ 700 SN Ia from the ESSENCE Project and other sources. They found no signature of shock heating from Red Giant companions. Bianco et al. (2011) determined that less than 10% out of 87 SN Ia could have come from white dwarf-Red Giant binary systems. Olling et al. (2015) used high cadence data from *Kepler* and they found no evidence for interaction in three SN Ia. The *Kepler* bandpass is not sensitive below 400 nm, and it is not clear if it would detect a shocked companion.

Individual SN Ia that are found very nearby or very early provide high quality data that may be unavailable to larger surveys. Analyses of such SN Ia have revealed an interesting amount of diversity, but no clear signs of shock interaction with a companion.

Brown et al. (2009, 2012) reported early *UV* LC from a few SN Ia without discovering evidence for excess luminosity. Foley et al. (2012a) described excess *UV* flux from SN 2009ig at early phases, but concluded that the colors were inconsistent with an interaction. Margutti et al. (2012) with X-rays and Chomiuk et al. (2012) with radio observations ruled out most

of the parameter space for a main-sequence or evolved companion to SN 2011fe. Schaefer & Pagnotta (2012) determined that SNR 0509–67.5 in the Large Magellanic Cloud contains no candidates for the companion star to deep limits. They claim to eliminate all previously published single-degenerate models for this SN Ia. Zheng et al. (2013) showed that measurements of SN 2013dy, obtained only a few hours after the estimated time of the explosion, do not reveal any evidence for interaction. Goobar et al. (2015) find that a very early section of the SN 2014J LC (~ 0.5 – 2.0 days after explosion) is flatter than a t^2 rise; but they rule out the SD model due to constraints on the size of the companion. Margutti et al. (2014) use X-ray non-detections of SN 2014J to rule out SD systems with steady mass loss.

Cao et al. (2015) report the detection of a significant *UV* pulse in early data of iPTF14atg. While intriguing, these results may not be directly applicable to the progenitor scenarios of normal SN Ia. This object is part of a class that is very rare (Ganeshalingam et al. 2012); it is about 3 magnitudes subluminous compared to a normal SN Ia, and it does not follow the Phillips relation (Phillips 1993). The current work focuses on the normal SN Ia, SN 2012cg. This object peaks in M_B between -19.4 and -19.8 mag and does follow the Phillips relation. Here we show that the early *UV* and optical LCs for SN 2012cg are well matched by SD models in which excess blue light is produced by the impact of the supernova on its companion. Photometric and spectroscopic observations of SN 2012cg and the data reduction details are described in Section 2. Excess luminosity in early photometry and spectra of SN 2012cg is described and analyzed in Section 3. Predictions of theoretical models for the interaction of a SN Ia and a binary companion are described in Section 4 and the models are compared to observations of SN 2012cg. Section 5 discusses the spectra and the evolution of features in pre-maximum spectra. Discussion and conclusions are presented in Section 6.

2. DATA ACQUISITION AND REDUCTION

SN 2012cg was discovered in Virgo Cluster Galaxy NGC 4424 on 2012 May 17.2, UT = MJD 56065.2 by the Lick Observatory Supernova Search (LOSS; Filippenko et al. 2001) with the 0.76 m Katzman Automatic Imaging Telescope (KAIT). The discovery was promptly announced to the community by email and by ATEL (Cenko et al. 2012). This rapid notice allowed many observers to begin following SN 2012cg on May 18 which was less than three days after the explosion and more than 16 days before the time of maximum brightness in the *B*-band ($t(B_{\max})$). We report photometric and spectroscopic observations of SN Ia 2012cg obtained from 2012 May 18.2, (UT) which is 16.1 days before $t(B_{\max})$ (-16.1 days) to June 26.0 ($+22.7$ days).

2.1. Photometric Observations

Figure 1 shows optical and near-infrared (NIR) photometry of SN 2012cg obtained from -16.1 to $+19.9$ days. The optical data are from the F. L. Whipple Observatory (FLWO) using the 1.2 m telescope and KeplerCam, while the NIR data were obtained with the Peters Automated Infrared Imaging Telescope (PAIRITEL).

The data in the figure have been corrected for Milky Way (MW) and host extinction. Polynomial fits give the dates of peak

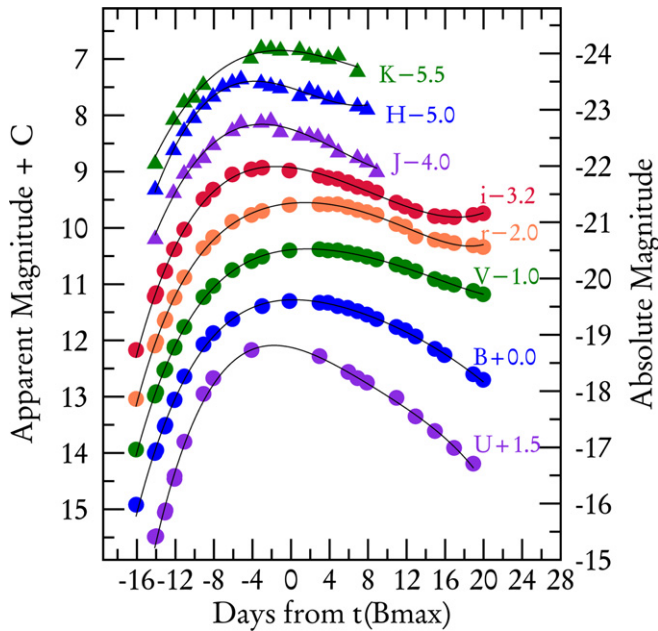


Figure 1. $UBVr'i'JHK_s$ photometry of SN 2012cg obtained at the F. L. Whipple Observatory from -16.1 to $+19.9$ days. The data are corrected for MW and host galaxy extinction. The black lines are polynomial fits used to identify the peak brightness and the time of peak for each filter (Table 1). We find $t(B_{\max}) = 560681.3$ (MJD) = June 3.3 (UT), with a peak magnitude in the B -band of -19.62 mag and $\Delta m_{15}(B) = 0.86$.

brightness in each filter, the maximum apparent magnitudes and the decline rate parameters, Δm_{15} (Table 1). The absolute magnitudes were computed using $d = 15.2 \pm 1.9$ Mpc ($\mu = 30.90 \pm 0.3$ mag; Tully–Fisher) for NGC 4424 (Cortes et al. 2008). Uncertainties in the absolute magnitude estimates do not affect our primary results.

The FLWO optical data (u' , B , V , r' , i') were reduced using IRAF and IDL procedures described in Hicken et al. (2007). Table 2 provides the original measurements without dereddening. The FLWO galaxy templates were obtained on 2014 January 10, which is 596 days after $t(B_{\max})$.

Transformation to the standard photometric system was performed using local comparison stars around the SN in the same field of view. The linear transformation equations were calibrated using Landolt (1992) standards for UBV and Smith et al. (2002) standards for r' - and i' -bands. The zero-points of the transformations were determined with data from photometric nights. The zero-points for images obtained on non-photometric nights were determined by differential aperture photometry (DAOPHOT) using tertiary standard stars in the vicinity of the SN. Further details of these methods can be found in Hicken et al. (2012).

The FLWO u' -band can be correlated with Landolt (1992) U -band magnitudes via the equation $u' = U + 0.854$ mag (Chonis & Gaskell 2008). Table 2 shows the measured values for u' while the figures use the values corrected for extinction and converted to U -band.

NIR images were obtained at the FLWO in the J , H , K_s bands by PAIRITEL (Table 3). The data are processed into mosaics using the PAIRITEL Mosaic Pipeline version 3.6 implemented in python. Photometry is performed on the mosaicked images with DoPHOT (Schechter et al. 1993) using a modified version of the ESSENCE project photometry pipeline (Miknaitis et al. 2007). Photometric zero-points are

Table 1
Peak Magnitudes for SN 2012cg FROM EXTINCTION-CORRECTED FLWO DATA

Band	MJD-56000 ^a	m_{λ} ^b	M_{λ} ^c	Δm_{15} ^d
KepCam U	79.6	10.59	-20.30	1.51
KepCam B	81.3	11.26	-19.63	0.86
KepCam V	82.3	11.36	-19.54	0.51
KepCam r'	81.8	11.53	-19.37	0.63
KepCam i'	79.5	12.08	-18.82	0.86
PAIRITEL J	78.3	12.16	-18.74	...
PAIRITEL H	77.6	12.39	-18.51	...
PAIRITEL K	80.2	12.35	-18.55	...

Notes.

^a ± 0.4 days.

^b ± 0.04 mag.

^c ± 0.08 mag.

^d ± 0.04 mag.

computed using the 2MASS point source catalog (Cutri et al. 2003). Details of PAIRITEL observations and reduction of NIR supernova data can be found in Friedman (2012) and Friedman et al. (2015).

We are fortunate to be able to include pre-maximum photometry of SN 2012cg that was obtained at other facilities. After their discovery of SN 2012cg, LOSS/KAIT continued to monitor the SN and their pre-maximum, uncorrected data are listed in Table 4. These data were reduced with an image-reduction pipeline described in Ganeshalingam et al. (2010). The acquisition and reduction of LOSS/KAIT data is described in detail by Silverman et al. (2012).

KAIT templates were obtained on 2013 December 9, which is 554 days after $t(B_{\max})$. The galaxy-subtracted photometry produced the same values as reported by Silverman et al. (2012) (W. Zheng 2016, private communication).

B - and V -band photometry was obtained by the Las Cumbres Observatory Global Telescope Network of 1 m telescopes (LCOGT), and reduced using a custom pipeline developed by LCOGT which is based on standard procedures, including *pyraf*, *DAOPHOT*, and *SWARP* in a python framework. Instrumental magnitudes were transformed to the standard system (Landolt 1992) using standard star observations obtained on photometric nights (Brown et al. 2013). The pre-maximum B - and V -band data from the LCOGT are listed in Table 5.

Swift photometric data were obtained with the Ultraviolet/Optical Telescope (UVOT; Roming et al. 2005). The UVOT reduction used a pipeline developed for the Swift Optical/Ultraviolet Supernova Archive (SOUSA; Brown et al. 2014). It is based on Brown et al. (2009), including subtraction of the host galaxy count rates and uses the revised UV zero-points and time-dependent sensitivity from Breeveld et al. (2011). Table 6 shows the v , b , u , $uvw1$, $uvw2$, and $uvw2$ measurements of SN 2012cg from -15.7 to $+1.0$ days. The *Swift* galaxy template was obtained by observing the location of SN 2012cg 428 days after $t(B_{\max})$.

ROTSE-IIIb uses an unfiltered CCD and an automated image differencing analysis to search for SNe (Yuan & Akerlof 2008). Photometry is calibrated to an effective r -band magnitude by comparing to USNO B1.0 (Monet et al. 2003). For comparison with other data sources in this paper, the

Table 2
FLWO 1.2 m Observations of SN 2012cg

MJD-56000	Phase ^a	Mag	Err
<i>u'</i> -band			
67.16	-14.1	15.84	0.03
67.31	-14.0	15.83	0.02
68.20	-13.1	15.41	0.02
68.27	-13.0	15.37	0.02
69.20	-12.1	14.81	0.02
69.21	-12.1	14.76	0.02
70.23	-11.1	14.15	0.02
72.22	-9.1	13.30	0.02
73.23	-8.1	13.02	0.02
77.20	-4.1	12.52	0.02
84.26	3.0	12.63	0.02
87.25	5.9	12.91	0.02
88.20	6.9	13.02	0.02
89.18	7.9	13.10	0.02
92.25	10.9	13.37	0.02
94.23	12.9	13.70	0.02
96.24	14.9	13.96	0.04
98.23	16.9	14.26	0.03
100.22	18.9	15.54	0.04
<i>B</i> -band			
65.22	-16.1	15.74	0.02
67.16	-14.1	14.82	0.02
67.31	-14.0	14.77	0.02
68.19	-13.1	14.35	0.02
68.26	-13.0	14.32	0.02
69.20	-12.1	13.88	0.02
69.21	-12.1	13.87	0.02
70.23	-11.1	13.46	0.02
72.21	-9.1	12.89	0.02
73.23	-8.1	12.69	0.02
75.21	-6.1	12.44	0.03
78.31	-3.0	12.21	0.02
81.14	-0.2	12.12	0.02
84.26	3.0	12.15	0.02
85.14	3.8	12.15	0.02
86.14	4.8	12.21	0.02
87.18	5.9	12.24	0.02
87.25	5.9	12.25	0.02
88.20	6.9	12.30	0.02
89.18	7.9	12.36	0.02
90.18	8.9	12.44	0.02
92.24	10.9	12.58	0.03
93.15	11.8	12.64	0.02
94.23	12.9	12.75	0.02
96.24	14.9	12.97	0.03
97.23	15.9	13.08	0.02
100.22	18.9	13.42	0.02
101.25	19.9	13.52	0.02
<i>V</i> -band			
65.21	-16.1	15.56	0.01
67.16	-14.1	14.60	0.01
67.30	-14.0	14.54	0.01
68.15	-13.2	14.15	0.01
68.26	-13.0	14.13	0.01
69.20	-12.1	13.75	0.01
69.21	-12.1	13.73	0.01
70.22	-11.1	13.38	0.01
72.21	-9.1	12.85	0.01
73.23	-8.1	12.65	0.01
75.21	-6.1	12.37	0.01

Table 2
(Continued)

MJD-56000	Phase ^a	Mag	Err
77.19	-4.1	12.21	0.01
78.30	-3.0	12.13	0.03
81.14	-0.2	12.02	0.01
84.25	2.9	12.00	0.01
85.14	3.8	12.02	0.01
86.14	4.8	12.02	0.02
87.17	5.9	12.06	0.01
87.25	5.9	12.06	0.01
88.20	6.9	12.09	0.01
89.18	7.9	12.13	0.01
90.18	8.9	12.18	0.01
92.24	10.9	12.27	0.01
93.15	11.8	12.32	0.01
94.23	12.9	12.39	0.01
96.23	14.9	12.53	0.01
97.23	15.9	12.59	0.01
98.23	16.9	12.63	0.01
100.22	18.9	12.74	0.01
101.25	19.9	12.80	0.01
<i>r'</i> -band			
65.21	-16.1	15.54	0.01
67.15	-14.2	14.59	0.01
67.30	-14.0	14.52	0.01
68.19	-13.1	14.15	0.01
68.26	-13.0	14.12	0.01
69.19	-12.1	13.74	0.01
69.21	-12.1	13.73	0.01
70.22	-11.1	13.38	0.01
72.21	-9.1	12.86	0.01
73.22	-8.1	12.67	0.01
75.21	-6.1	12.39	0.01
77.19	-4.1	12.27	0.01
78.30	-3.0	12.20	0.01
81.14	-0.2	12.09	0.01
84.25	2.9	12.08	0.01
85.14	3.8	12.08	0.01
86.14	4.8	12.08	0.01
87.17	5.9	12.13	0.01
87.24	5.9	12.13	0.01
88.20	6.9	12.18	0.01
89.18	7.9	12.22	0.01
90.18	8.9	12.27	0.01
92.24	10.9	12.43	0.01
93.14	11.8	12.48	0.01
94.23	12.9	12.65	0.09
96.23	14.9	12.71	0.01
97.23	15.9	12.73	0.01
98.23	16.9	12.77	0.01
100.22	18.9	12.81	0.01
101.25	19.9	12.84	0.01
<i>i'</i> -band			
65.21	-16.1	15.71	0.01
67.15	-14.2	14.76	0.01
67.30	-14.0	14.70	0.01
68.18	-13.1	14.30	0.07
68.25	-13.1	14.31	0.01
69.14	-12.2	13.92	0.01
69.20	-12.1	13.92	0.01
70.22	-11.1	13.57	0.01
72.21	-9.1	13.03	0.01
73.22	-8.1	12.86	0.01
75.21	-6.1	12.59	0.01

Table 2
(Continued)

MJD-56000	Phase ^a	Mag	Err
77.19	-4.1	12.50	0.01
78.30	-3.0	12.47	0.03
81.14	-0.2	12.52	0.01
84.25	2.9	12.61	0.01
85.14	3.8	12.65	0.01
86.14	4.8	12.67	0.01
87.17	5.9	12.75	0.01
87.24	5.9	12.73	0.01
88.20	6.9	12.81	0.01
89.17	7.9	12.84	0.01
90.17	8.9	12.91	0.01
92.24	10.9	13.09	0.01
93.14	11.8	13.16	0.01
94.22	12.9	13.24	0.01
96.23	14.9	13.33	0.01
97.23	15.9	13.34	0.01
98.23	16.9	13.35	0.01
100.22	18.9	13.31	0.01
101.25	19.9	13.28	0.01

Note.^a Estimated date of $t(B_{\max})$: MJD 56081.3 = June 3.3 (UT).

ROTSE clear data, which include significant B - and U -band sensitivity, have their zero-point adjusted to B -band. ROTSE detected SN2012cg on May 17.178, which is 1.1 hr before the discovery epoch reported by Silverman et al. (2012) (May 17.223). Prediscovery ROTSE images on May 16.177 yielded no detection to a limiting magnitude of 16.9. Table 7 gives the ROTSE measurements.

The ROTSE data are reduced by differential aperture photometry using IDL procedures adapted from DAOPHOT, followed by subtraction of the underlying host. Transformation to Bessel V -band was performed using reference stars to a radius of $3'$ from SN 2012cg. The transformation was calibrated by comparison to the APASS catalog. This measurement is compared to image differencing where it exists. There is good agreement with the host subtraction result except for the first epoch when the SN is dim and image differencing is difficult. Additional uncertainties at all epochs are extracted from the measured variation due to altering the aperture photometry parameters.

2.2. Reddening

To estimate reddening due to the host galaxy, we compared the $B - V$ color curves of SN 2012cg to SN 2011fe after applying $E(B - V)_{\text{MW}} = 0.018$ mag for the Milky Way extinction in the direction of SN 2012cg (Schlafly & Finkbeiner 2011). B - and V -band LC of SN 2011fe were obtained from the Piszkestető Mountain Station of the Konkoly Observatory, Hungary. These data were previously published by Vinkó et al. (2012).

The color curves are aligned at $t(B_{\max})$ by applying a vertical shift of $E(B - V) = 0.18$ mag. SN 2011fe is essentially unreddened by its host (Nugent et al. 2011; Vinkó et al. 2012; Chomiuk 2013), so this difference provides a plausible estimate of the reddening of SN 2012cg from the host galaxy. It is the

Table 3
PAIRITEL Observations of SN 2012cg

MJD-56000	Phase ^a	Mag	Err
<i>J</i> -band			
67.2	-14.1	14.36	0.05
69.1	-12.2	13.54	0.06
70.2	-11.1	13.20	0.08
71.2	-10.1	13.01	0.02
72.2	-9.1	12.92	0.09
73.2	-8.1	12.69	0.02
75.2	-6.1	12.44	0.02
76.1	-5.2	12.30	0.08
78.2	-3.1	12.29	0.04
79.2	-2.1	12.26	0.08
80.2	-1.1	12.46	0.10
82.2	0.9	12.52	0.07
83.2	1.9	12.52	0.02
84.1	2.8	12.55	0.04
85.2	3.9	12.65	0.03
86.2	4.9	12.82	0.06
88.2	6.9	12.92	0.02
89.2	7.9	13.00	0.03
90.2	8.9	13.17	1.13
<i>H</i> -band			
67.2	-14.1	14.42	0.11
69.1	-12.2	13.72	0.02
70.2	-11.1	13.38	0.02
71.2	-10.1	13.15	0.03
72.2	-9.1	12.91	0.02
73.2	-8.1	12.77	0.02
74.2	-7.1	12.59	0.06
75.2	-6.1	12.52	0.06
76.1	-5.2	12.46	0.08
78.2	-3.1	12.53	0.02
79.2	-2.1	12.58	0.04
80.2	-1.1	12.62	0.05
82.2	0.9	12.76	0.09
83.2	1.9	12.65	0.07
84.1	2.8	12.72	0.03
85.2	3.9	12.82	0.02
86.2	4.9	12.82	0.04
88.2	6.9	12.94	0.02
89.2	7.9	13.00	1.52
<i>Ks</i> -band			
67.2	-14.1	14.42	0.11
69.1	-12.2	13.64	0.08
70.2	-11.1	13.33	0.09
71.2	-10.1	13.25	0.04
72.2	-9.1	13.02	0.02
77.1	-4.2	12.55	0.07
78.2	-3.1	12.37	0.07
79.2	-2.1	12.38	0.04
80.2	-1.1	12.41	0.02
82.2	0.9	12.40	0.03
83.2	1.9	12.50	0.04
84.1	2.8	12.53	0.03
85.2	3.9	12.56	0.02
86.2	4.9	12.50	0.08
88.2	6.9	12.79	0.10

Note.^a Estimated date of $t(B_{\max})$: MJD 56081.3 = June 3.3.

Table 4
KAIT (Premaximum) Observations of SN 2012cg

MJD-56000	Phase ^a	Mag	Err
<i>B</i> -band			
64.2	-17.1	17.28	0.09
65.2	-16.1	15.94	0.05
66.2	-15.1	15.24	0.03
67.2	-14.1	14.83	0.03
68.2	-13.1	14.33	0.03
69.2	-12.1	13.90	0.03
70.2	-11.1	13.50	0.03
71.2	-10.1	13.11	0.03
73.2	-8.1	12.67	0.03
75.2	-6.1	12.36	0.03
76.2	-5.1	12.30	0.03
78.2	-3.1	12.11	0.03
79.2	-2.1	12.09	0.03
81.2	-0.1	12.12	0.03
<i>V</i> -band			
64.2	-17.1	17.15	0.07
65.2	-16.1	15.56	0.02
66.2	-15.1	14.96	0.02
67.2	-14.1	14.58	0.02
68.2	-13.1	14.08	0.02
69.2	-12.1	13.74	0.02
70.2	-11.1	13.36	0.02
71.2	-10.1	13.08	0.02
73.2	-8.1	12.62	0.02
75.2	-6.1	12.36	0.02
76.2	-5.1	12.28	0.02
78.2	-3.1	12.08	0.02
79.2	-2.1	12.12	0.02
81.2	-0.1	12.04	0.02

Note.

^a Estimated date of $t(B_{\max})$: MJD 56081.3 = June 3.3.

Table 5
LCOGT (Premaximum) Observations of SN 2012cg

MJD-56000	Phase ^a	Mag	Err
<i>B</i> -band			
66.6	-14.7	15.23	0.10
69.6	-11.7	13.87	0.09
70.8	-10.5	13.35	0.06
75.6	-5.7	12.40	0.08
83.6	2.3	12.02	0.05
<i>V</i> -band			
66.6	-14.7	15.00	0.07
69.6	-11.7	13.72	0.06
70.8	-10.5	13.33	0.04
75.6	-5.7	12.40	0.06
83.6	2.3	11.97	0.05

Note.

^a Estimated date of $t(B_{\max})$: MJD 56081.3 = June 3.3.

same $B - V$ offset found by Silverman et al. (2012) for SN 2012cg. Figure 2 illustrates the comparison between the two de-reddened color curves. The dotted lines mark the ± 0.05 mag uncertainty of $E(B - V)_{\text{host}}$ for SN 2012cg.

Table 6
Swift (Premaximum) Observations of SN 2012cg

MJD-56000	Phase ^a	Mag	Err
<i>U</i> -band			
65.8	-15.5	15.46	0.07
67.2	-14.1	15.13	0.07
71.4	-9.9	12.91	0.05
73.5	-7.8	12.20	0.03
80.7	-0.6	11.72	0.03
82.4	1.1	11.78	0.03
<i>B</i> -band			
65.8	-15.5	15.48	0.06
67.2	-14.1	14.93	0.05
71.4	-9.9	13.25	0.04
73.5	-7.8	12.77	0.04
80.7	-0.6	12.16	0.03
82.4	1.1	12.19	0.03
<i>V</i> -band			
65.8	-15.5	15.23	0.09
67.2	-14.1	14.62	0.06
71.4	-9.9	13.09	0.04
73.5	-7.8	12.61	0.04
75.2	-6.1	12.42	0.03
78.9	-2.4	12.14	0.03
80.6	-0.7	12.08	0.03
82.5	1.2	12.04	0.03
<i>UVW1</i> -band			
65.8	-15.5	17.14	0.12
67.6	-13.7	16.89	0.11
68.4	-12.9	16.46	0.18
71.4	-9.9	14.60	0.05
73.3	-8.0	13.94	0.04
75.2	-6.1	13.57	0.04
78.9	-2.4	13.33	0.04
80.6	-0.7	13.39	0.04
82.5	1.2	13.48	0.03
<i>UVM2</i> -band			
65.8	-15.5	19.82	0.35
67.2	-14.1	19.86	0.32
70.4	-10.9	18.55	0.25
71.4	-9.9	18.02	0.12
74.3	-7.0	17.04	0.11
75.2	-6.1	16.76	0.10
78.9	-2.4	16.50	0.10
80.6	-0.7	16.49	0.09
82.5	1.2	16.48	0.09
<i>UVW2</i> -band			
65.8	-15.5	18.83	0.18
67.2	-14.1	18.43	0.15
70.5	-10.8	16.79	0.11
71.4	-9.9	16.34	0.11
73.9	-7.4	15.73	0.10
74.4	-6.9	15.51	0.09
75.2	-6.1	15.40	0.07
78.9	-2.4	15.14	0.08
80.6	-0.7	15.18	0.07
82.5	1.2	15.19	0.06

Note.

^a Estimated date of $t(B_{\max})$: MJD 56081.3 = June 3.3.

Table 7
ROTSE (Premaximum) Observations of SN 2012cg

MJD-56000	Phase ^a	Mag	Err
Clear Filter			
64.2	-17.1	16.63	0.38
65.2	-16.1	15.41	0.11
66.2	-15.1	15.39	0.45
67.2	-14.1	14.88	0.03
68.2	-13.1	14.32	0.03
70.2	-11.1	13.37	0.20
72.2	-9.1	12.95	0.03
73.2	-8.1	12.70	0.03
75.2	-6.1	12.55	0.02
76.2	-5.1	12.43	0.02

Note.

^a Estimated date of $t(B_{\max})$: MJD 56081.3 = June 3.3.

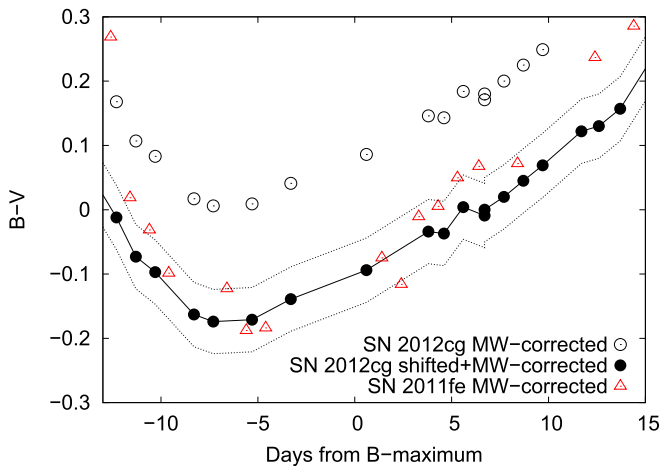


Figure 2. $B - V$ color evolution for SN 2012cg and SN 2011fe after correction for Milky Way (Galactic) reddening. The curves are fit to minimize the difference near $t(B_{\max})$. We assume that SN 2011fe has essentially zero reddening, so the magnitude of the difference determines $E(B - V)_{\text{host}} \approx 0.18$ mag. (See Section 2.)

Summing the MW and host galaxy components of the extinction, we determine that the total reddening of SN 2012cg is $E(B - V)_{\text{total}} = 0.18 + 0.018 = 0.198 \approx 0.20 \pm 0.05$ mag, which is consistent with that reported by Silverman et al. (2012).

We also use the hierarchical Bayesian statistical model BayeSN to fit the $BVr'i'JH$ LC data from FWLO and PAIRITEL. BayeSN models the variations in observed optical and NIR SN Ia LCs as a combination of an intrinsic LC distribution and a distribution of host galaxy dust extinction, determined from a low- z SN Ia training set (Mandel et al. 2011). Applied to an individual SN Ia, it computes the posterior probability of its LC and dust parameters. Using this method, for SN 2012cg, we inferred a host galaxy dust color excess of $E(B - V)_{\text{host}} = 0.25 \pm 0.03$ (assuming $R_V = 3.1$).

Differences in the total reddening on the order of 0.08 mag have much smaller effect on the $B - V$ curve of SN 2012cg than the color excess we measure (see Section 3). In Section 4.1 we show that the color excess for SN 2012cg with respect to other SNe Ia is about 0.4–0.5 mag, which is significantly

larger than the uncertainty of $E(B - V)$ given above. We can increase the uncertainties to ± 0.08 mag include the BayeSN result and the Amanullah et al. (2015) result of $E(B - V)_{\text{total}} = 0.15 \pm 0.02$ mag without affecting our primary results.

2.3. Estimating $t(B_{\max})$ and $t = 0$

We estimate $t(B_{\max})$ by fitting a polynomial to the B -band LC (Figure 1). We find $t(B_{\max}) = \text{MJD } 56081.3 \pm 0.5$ days (June 3.3 UT), with $M_B = -19.62 \pm 0.08$ mag (Table 1). The phases of all reported observations of SN 2012cg are defined as relative to the time of $t(B_{\max})$. Our estimate for $t(B_{\max})$ is between the estimates provided by Silverman et al. (2012) (June 2.0 UT) and Munari et al. (2013) (June 5.0 UT).

We also estimate the rise time from explosion ($t = 0$) to $t(B_{\max})$ of SN 2012cg to compare the timing of observations with other SN. This estimate has no effect on our measurements of excess flux. The $L \propto t^2$ “fireball” model for early LCs of SN Ia (Arnett 1982) produces a good fit to the data for many well observed SN (Nugent et al. 2011). The Arnett (1982) model derives $L \propto t^2$ from solving a diffusion equation, but the result is much like the simple assumption that with a constant temperature and expansion velocity, flux scales with the surface area which is proportional to t^2 . We note however, that recent studies have shown that the LC for some SN Ia are not well fit by the t^2 model (e.g., Piro & Nakar 2013). Note also that the t^2 model is strictly valid only for the bolometric LC. Although t^2 works for optical band passes as well in some circumstances, in general for filtered LCs one may expect deviations from the strict $n = 2$ exponent. It seems reasonable to use a more general t^n model having different n indices for different filter bands. We find the n does vary among the bands in our fits (Section 3).

Zheng et al. (2013) fit the LC SN 2013dy with a variable, or “broken,” power law that has a very steep rise (a higher power law index) for the first day. At a little more than one day after first light, this model adopts a more gradual curve with an exponent of 2.24 ± 0.08 . Zheng et al. (2014) find that a similar model is required to fit the LC of SN 2014J. Dessart et al. (2014) present LC for delayed detonation and pulsational-delayed detonation models of SN Ia that do not fit well to the t^2 model.

Despite differences in their trajectories back toward $t = 0$, if we ignore the breakout phase, then all of these models are monotonically decreasing without inflections in the model LC. At this time, it is not obvious that one model is preferred to another, so we use the classic t^2 model to find the moment of explosion. The model LC is fit to the FLWO data between -12 and -8 days which produces an estimated time for the explosion of $\text{MJD} = 56062.5 \pm 0.5$ days (May 15.5 UT). This result is consistent with Silverman et al. (2012) who estimate the explosion date to be May 15.7 (UT).

These results give a rise time of 18.8 days. A variable power law model (Zheng et al. 2013, 2014) would estimate an explosion date of about one day later and a rise time about one day faster. Silverman et al. (2012) measure an earlier time of $t(B_{\max})$ (June 2.0 UT), so their rise time estimate is shorter at 17.3 days.

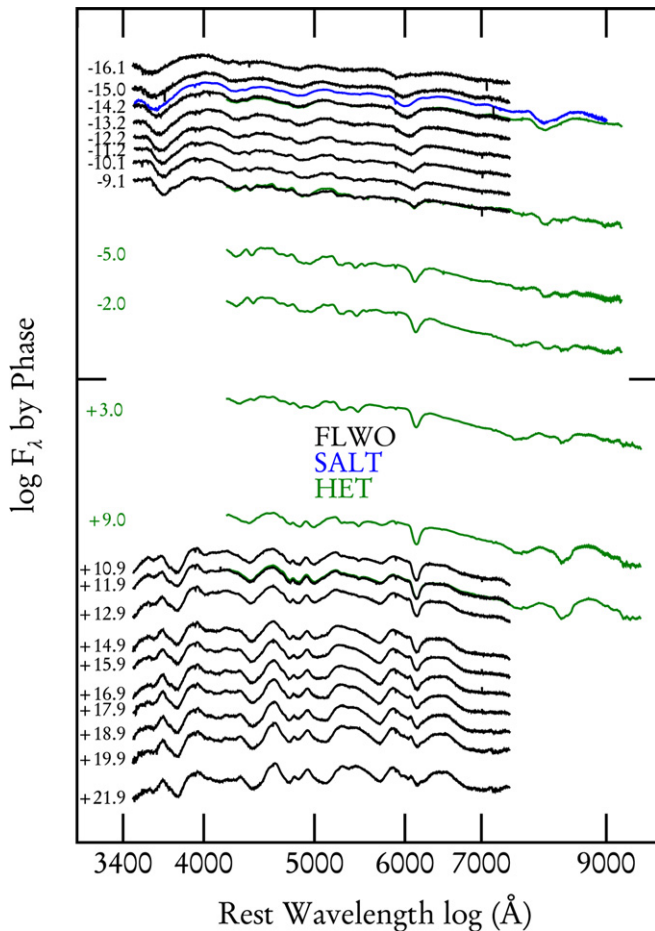


Figure 3. Optical spectroscopy of SN 2012cg obtained from -16.1 to $+21.9$ days. The sources are FLWO (black), SALT (blue), and HET (green). The continuum slopes of the spectra show that SN 2012cg is a blue SN Ia at early phases. Absorption features in the pre-maximum spectra are relatively weak. This is characteristic of slightly overluminous SN Ia.

2.4. Zero-point Corrections

To calibrate and compare the early-time data of SN 2012cg from different sources, a $(t - t_{\text{exp}})^2$ model LC was fit to the V - and B -band data. The model was selected by minimizing the residuals for the FLWO data at phases from -14 to -10 days. Each of the other data sets was fit to the t^2 curve by moving it up or down to minimize the residuals of the data from each source over the same phase interval. Thus the shape of each LC was preserved while systematic errors in the photometric calibration of the LC from the various sources was reduced.

The zero-point offsets used for all the LC figures presented herein are: 0.00 mag for KAIT V -band; -0.08 mag for KAIT B -band; -0.12 mag for LCOGT V -band; -0.24 mag for LCOGT B -band; -0.16 mag for *Swift* V -band; -0.08 mag for *Swift* B -band. The ROTSE data, has its zero-point adjusted to B -band.

We note that the -17 days KAIT data in both V - and B -bands (blue triangles) were measured from data near the detection limit of the instrument; the error bars come from the formal measurement uncertainty and are thus likely underestimated (Silverman et al. 2012).

²² Upon acceptance of this work, all spectra herein will be available in WISeREP (the Weizmann Interactive Supernova data REpository; Yaron & Gal-Yam 2012): <http://wiserep.weizmann.ac.il/>.

Table 8
Optical Spectra of SN 2012cg

Date	Phase ^a	Telescope/Instrument	Range (Å)	R ($\lambda/\Delta\lambda$)
May 18.21	-16.1	FLWO/FAST	3300–7400	2700
May 19.30	-15.0	FLWO/FAST	3300–7400	2700
May 19.81	-14.5	SALT/RSS	3500–9000	1100
May 20.15	-14.2	FLWO/FAST	3300–7400	2700
May 20.20	-14.1	HET/LRS	4100–10200	1500
May 21.15	-13.2	FLWO/FAST	3300–7400	2700
May 22.15	-12.2	FLWO/FAST	3300–7400	2700
May 23.15	-11.2	FLWO/FAST	3300–7400	2700
May 24.16	-10.1	FLWO/FAST	3300–7400	2700
May 25.16	-9.1	FLWO/FAST	3300–7400	2700
May 25.20	-9.0	HET/LRS	4100–10200	1500
May 29.20	-5.0	HET/LRS	4100–10200	1500
June 1.20	-2.0	HET/LRS	4100–10200	1500
June 6.20	$+3.0$	HET/LRS	4100–10200	1500
June 12.20	$+9.0$	HET/LRS	4100–10200	1500
June 14.21	$+10.9$	FLWO/FAST	3300–7400	2700
June 15.15	$+11.9$	FLWO/FAST	3300–7400	2700
June 15.20	$+12.0$	HET/LRS	4100–10200	1500
June 16.25	$+12.9$	FLWO/FAST	3300–7400	2700
June 18.15	$+14.9$	FLWO/FAST	3300–7400	2700
June 19.19	$+15.9$	FLWO/FAST	3300–7400	2700
June 20.24	$+16.9$	FLWO/FAST	3300–7400	2700
June 21.22	$+17.9$	FLWO/FAST	3300–7400	2700
June 22.15	$+18.9$	FLWO/FAST	3300–7400	2700
June 23.18	$+19.9$	FLWO/FAST	3300–7400	2700
June 25.17	$+21.9$	FLWO/FAST	3300–7400	2700

Note.

^a Phase in days with respect to the estimated date of $t(B_{\text{max}})$: MJD 56081.3 = June 3.3.

2.5. Spectroscopic Observations

Optical spectra²² of SN 2012cg were obtained from May 18.2 through June 25.2 (Figure 3). These dates correspond to phases -16.1 to $+21.9$ days. Details for the observations are in Table 8.

Optical spectra (3480–7420 Å, displayed in black) were obtained with the FLWO 1.5 m Tillinghast telescope and the FAST spectrograph (FAST; Fabricant et al. 1998). The position angle was 90° but the airmass was low (≤ 1.18). FAST data are reduced using a combination of standard IRAF and custom IDL procedures (Matheson et al. 2005). Additional optical spectra (4,200–10,100 Å, green) were obtained with the 9.2 m Hobby–Eberly Telescope (HET; Ramsey et al. 1998) at the McDonald Observatory using the Marcario Low-Resolution Spectrograph (LRS; Hill et al. 1998). HET/LRS spectra are reduced with standard IRAF procedures. The HET spectra of SN 2012cg obtained on May 20, 25, and 29 were previously published by Silverman et al. (2012).

One early epoch spectrum was obtained with Robert Stobie Spectrograph (RSS) on the Southern African Large Telescope (SALT) covering the range 3500–9000 Å (blue). This was reduced with a custom pipeline that incorporates routines from PyRAF and PySALT (Crawford et al. 2010).

Optical spectra from the FLWO, HET and SALT cover slightly different wavelength ranges. None of the sources provide continuous coverage through these phases (see Table 8). When they do overlap (-14 , -9 , 12 days), the agreement is excellent.

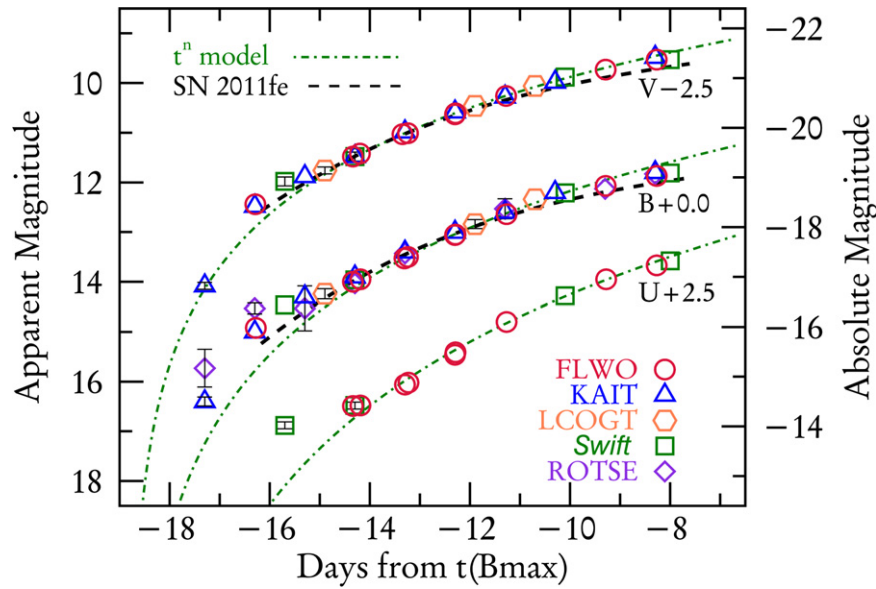


Figure 4. V -, B -, and U -band photometry (top to bottom) of SN 2012cg from multiple sources. The data have been corrected for extinction. The LC of SN 2011fe (Vinkó et al. 2012, dashed black lines) stretched to yield a rise time of 18.8 days equivalent to SN 2012cg (see the text) and t^n model LCs (dotted–dashed, green) are plotted as templates for a normal SN Ia. The power-law indices used for different bands are 3.4 (U), 2.4 (B), and 2.2 (V), see the text for details. In all bands, the SN 2012cg data display excess flux at phases earlier than -14 days. From -14 days toward maximum light, the data from SN 2012cg and SN 2011fe fit the templates well. Uncertainties are marked only where they approach the size of the symbols. Note that uncertainties on the earliest KAIT points are likely underestimated (see Section 4 for more information).

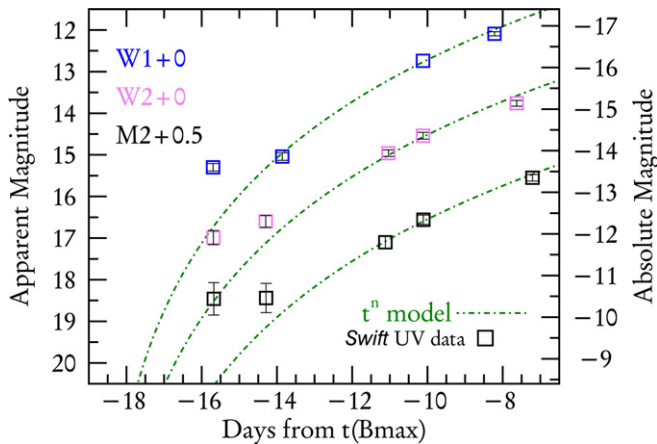


Figure 5. *Swift* photometry (squares) for SN 2012cg obtained in UV filters: $W1$, $M2$, and $W2$. The data have been corrected for extinction. Excess flux is apparent at -16 days in all filters, and it is also present at -14 days for $M2$ and $W2$. Model LCs scaled as $t^{3.6}$ (see text) are plotted for reference in each passband with green dotted–dashed lines. The timing of the observed UV excess matches well with the optical data displayed in Figure 4.

3. DETECTION OF EXCESS LUMINOSITY

Figures 4 and 5 display early photometry in six filters from five sources that reveal excess luminosity in the very early LC of SN 2012cg. These data have been corrected for extinction using $E(B - V)_{\text{total}} = 0.20$ mag (see Section 2.2) and the Fitzpatrick (1999) reddening law with $R_V = 3.1$. By “excess” we mean luminosity measurements that exceed the fit model function by at least three times the measurement uncertainty, i.e., they deviate from the models by more than 3σ . For each passband in Figures 4 and 5, a t^n model LC (green dotted–dashed curves) was fit to the measurements between -14 and -8 days. Regardless of whether n was fixed at 2 or allowed to

float, for phases earlier than -14 days the observed data from various sources consistently exceed the model LC by several times the measurement uncertainty.

By using strict t^2 models we find that between -18 and -14 days the average flux excess ratio, $(f_{\text{obs}} - f_{\text{model}})/(\sigma_{\text{obs}}^2 + \sigma_{\text{model}}^2)^{1/2}$, is 12.9, 4.9, and 4.5 for the U -, B - and V -band, respectively (note that the amount of excess flux is higher in the B -band than in the V -band, but the higher uncertainty of the B -band observations make their excess flux ratio similar to that of the V -band). Using the more general t^n model and letting n float, $n = 3.4 \pm 0.1$ (U), $n = 2.44 \pm 0.05$ (B), and $n = 2.20 \pm 0.03$ (V) are found when fitting the data between -14 and -8 days as before. In this case the excess flux ratios are somewhat reduced, but their average values are still 7.7 in U , 4.4 in B , and 3.9 in V . Thus, regardless of which models are used for comparison, the earliest observed fluxes deviate from the models by more than 4σ in the V -band and blueward. Note that a model LC that uses a variable power-law with the date of explosion scaled to one day later than the t^2 model and higher indices for the first day (Zheng et al. 2013, 2014) would increase the measured excess fluxes even more.

In order to test the null-hypothesis, i.e., whether the entire LC could be fit with a single t^n model without any excess flux, we extended the fitting range to -18 days, thus, all data in between -18 and -8 days were fit by the more flexible t^n model. This resulted in lower indices than in the previous case: $n = 3.2 \pm 0.1$, 2.19 ± 0.05 , and 2.05 ± 0.04 for U , B , and V , respectively. However, the quality of the fit drastically worsened, resulting in $\chi^2/\text{dof} = 33.2$ (U), 17.5 (B), and 11.1 (V) compared to 16.2, 4.9 and 2.3 given by the previous fitting. This fit obviously reduces the deviations of the earliest data from the model curve, and the worse quality of the fit increases the uncertainty of the model (σ_{model} above). Yet, in the data earlier than -15 days the average excess flux is still present in U (6.1σ) and in B (2.0σ), even though it is diminished in the

V -band (0.6σ). It is emphasized that the decrease of the signal-to-noise of the excess flux in this model fit is almost entirely due to the increase of σ_{model} (as a consequence of the much higher χ^2), compared to the previous cases when only the data taken after -14 days were fit. We conclude that for the B - and U -band the null-hypothesis can be rejected with high confidence, thus, the excess flux earlier than -15 days is inevitably present in our U - and B -band data. For the V -band data the excess flux is not statistically significant in the model which fits the entire LC, but becomes more evident in the models that are constrained by data taken later than -14 days, which also fit the LC much better.

We emphasize that we do not use these simple power-law LC models either to derive physical constraints for the amount of excess flux or to investigate their origin. The physics of the origin of the early-phase excess fluxes will be studied by comparing the data with real physical models in Section 4. Here we use the t'' models only for illustration purpose, i.e., to reveal that, unlike other well-observed SN Ia, SN 2012cg show noticeable deviations from the simple “fireball” model in the blue- and UV -bands.

Figure 4 shows measurements from V -, B -, and U -band filters. The abscissa is time in days from $t(B_{\text{max}})$, and the ordinate is apparent magnitude with absolute magnitude (using $\mu = 30.90$) displayed on the right axis. Uncertainties are smaller than the size of the symbols, excepted where indicated. The optical data were obtained at the FLWO, KAIT, LCOGT, *Swift*, and ROTSE.

SN 2011fe has been described as a good example of a typical SN Ia (e.g., Nugent et al. 2011; Chomiuk 2013). We use data of SN 2011fe as a surrogate for a “normal” SN Ia, using the well constrained LC parameters of Vinkó et al. (2012) and Pereira et al. (2013). In Figure 4, V - and B -band LC of SN 2011fe (Vinkó et al. 2012) are plotted as black dashed lines. The extinction-corrected data for SN 2011fe are fit with a polynomial and plotted as a continuous LC in order to avoid crowding the figure with more symbols. The rise time for SN 2011fe is 17.7 days, whereas our estimate for the rise time of SN 2012cg is 18.8 days. Since we make all direct comparisons based on the phase with respect to $t(B_{\text{max}})$, we stretch the rise time of SN 2011fe to 18.8 days for plotting.

Figure 5 shows *Swift* measurements for UV filters: $W1$, $M2$, and $W2$. Error bars in the figure and uncertainties in the table take into account the leaks of redder light into both the $W1$ and $W2$ filters. It is clear that the UV measurements at -16 days are well above the power-law model LCs. Using power-law indices similar to those found above for the U -band ($n \sim 3.6$) the excess fluxes for the *Swift* UV -bands are similar to that in the U -band. The larger error bars in the *Swift* data decrease the excess flux ratios, but they are still significant: 9.1 in $W1$, 3.0 in $M2$ and 4.9 in $W2$ filters, respectively. The amount of excess flux is greater in the UV at -16 days than it is for the B and V optical bands shown in Figure 4. This result is consistent with model predictions for increased UV luminosity from interaction events (see Section 4).

4. COMPARING SN 2012CG TO MODELS

In the SD model, assuming the absence of circumstellar material, a SN Ia will expand freely after the explosion until it encounters the companion star. At the point of impact, matter is compressed and heated while the SN continues to expand. Material flowing around the companion star forms a bow

shock, and a cavity is opened in the SN ejecta as it is diverted around the companion. Emission from the shock heated region can escape through this hole in the expanding SN (Marietta et al. 2000; Kasen 2010; Cao et al. 2015). The size of the cavity is determined by the radius of the companion and that determines the amount of excess radiation that will emerge within the first few hours.

The outer layers of the expanding material fill in the hole, but the supernova ejecta continue to collide with the binary companion. These impacting layers are heated by compression and some of the kinetic energy is also dissipated at this point. Additional thermal energy can diffuse out in the hours and days that follow the initial prompt burst. Continuing radiative diffusion from deeper layers of ejecta can produce emission in the optical and UV that may exceed the radioactively powered luminosity of the supernova for a few days after the explosion.

Kasen (2010) modeled the shock from a SN Ia as it impacts a companion star. He calculated the observational consequences for three different binary companions: a Red Giant with $r = 2 \times 10^{13}$ cm, a $6M_{\odot}$ MS star with $r = 2 \times 10^{12}$ cm, and a $2M_{\odot}$ MS star with $r = 5 \times 10^{11}$ cm. The models predict that under optimal conditions, optical and UV emission from the interaction will produce a detectable contribution to the LCs for a few days after the explosion. After this time, the shock heated emission will no longer contribute to the observed LC, and the SN will behave like a normal SN Ia with the LC entirely powered by radioactive decay of ^{56}Ni .

The reddest filter modeled by Kasen (2010) is V . Consequently we do not consider the r'/R and i'/I data when discussing the interaction models. The models were originally defined in AB magnitudes and the observations were measured in the Vega system. In order to compare them, we changed the magnitude measurements for both data and the models to flux units. In that format, the models can be moved up or down while preserving the relative flux at all wavelengths.

In Figure 6, we show the same U -, B -, and V -band data of SN 2012cg that appear in Figure 4. The data are plotted with model LC from for a normal SN Ia with no interaction (solid black line) and LC for interactions with three possible companions: a Red Giant star (RG, red), a $6M_{\odot}$ main-sequence star (MS, blue) and a $2M_{\odot}$ MS star (green) Kasen (2010). The models for different passbands show that the signatures of interaction have a greater deviation from normal at bluer wavelengths.

With the exception of the RG model that diverges from the other models at later phases, the interaction models remain close together and fit the data well from -14 to -8 days. Earlier than -14 days however, the predicted LC for the different models diverge and the separation is sufficient to easily differentiate the model predictions.

In Figure 6, the -16 and -15 days data are clearly brighter than the non-interaction models. There is some scatter in the data, but most of the points lie closest to the blue line which is the model for interaction with a $6M_{\odot}$ MS companion. The data are inconsistent with models for a normal SN Ia or a Red Giant companion viewed on axis.

This result defines a $6M_{\odot}$ MS star as the smallest allowed companion. If the impact location were not directly along the line of sight to the SN, then the observations of SN 2012cg could be produced by a much larger companion, such as a Red Giant. As the viewing angle becomes more oblique, the observed excess flux would be reduced.

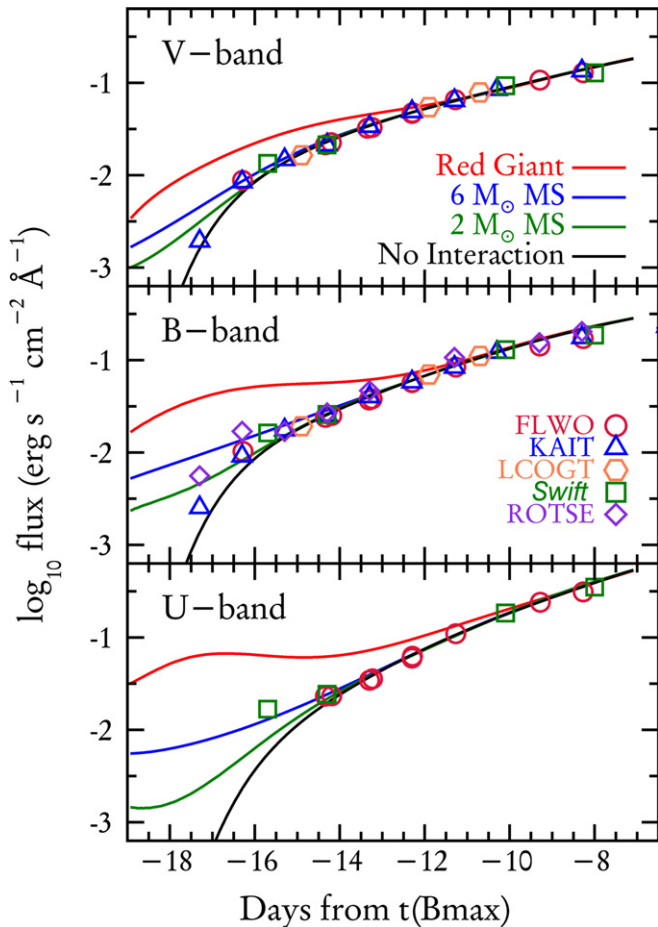


Figure 6. Early U -, B -, and V -band data for SN 2012cg are plotted with model results for a normal SN Ia with no interaction and 3 cases of interaction with a binary companion (Kasen 2010). The models fit the data well at phases after -14 days (except for the Red Giant model (red line) that is divergent through about -11 days). When the model light curves separate, the model for a SN Ia with no interaction (black) has the lowest predicted flux. Data obtained before -14 days are brighter than predicted by the normal models and the closest model fit is for interaction with a $6 M_{\odot}$ MS star (blue line). The values of the earliest KAIT points (blue triangles) are somewhat uncertain (see Section 4 for more information).

Off-center interaction sites will be more common than directly aligned sites. Kasen (2010) determined that even for the strongest possible signal from an interaction, the collision of a SN Ia with a Red Giant companion, the increased flux due to shock-heated ejecta would only be detectable about 10% of the time. Smaller companions create smaller and weaker interaction signatures that are even more dependent on viewing angle in order to be detected. Brown et al. (2012) found similar results for a decrease in the observed flux as the position of the interaction is incrementally offset from the direct line of sight to the SN. They also find that the observed flux as a function of wavelength changes, with the UV becoming relatively brighter as the viewing angle increases.

4.1. $B - V$ Colors

Figure 7 shows $B - V$ colors for SN 2012cg from -17 to -8 days. Blue colors are negative (toward the bottom of the figure) and red colors are positive (toward the top). The data and models are the same as those presented in the top 2 panels

of Figure 6. We note that the colors presented here match well with those reported by Silverman et al. (2012).

The Kasen (2010) model for a normal SN Ia without interaction (solid black line) shows $B - V$ colors that are significantly different than the $B - V$ colors for models with interaction. The non-interacting model is very red (top left) soon after the explosion, due to Fe-group line blending that suppresses the B -band continuum. The $B - V$ colors rapidly become bluer as the relative B -band flux increases. For a non-interacting SN Ia, $B - V$ is monotonically decreasing during this time period and it never produces a color peak.

The figure also shows $B - V$ colors for two other normal SN Ia that have very early observations. The filled stars are $B - V$ for SN 2009ig (Foley et al. 2012a) and the filled diamonds are $B - V$ for SN 2011fe (Pereira et al. 2013). In both cases, the color curves follow the general path of Kasen’s non-interacting model. They become bluer with time and neither of these normal SN Ia produces a peak in their color curves. The phases of both SN Ia have been stretched to match the 18.8 day rise time of SN 2012cg. Without stretching, the color curves rise even more steeply into the red.

The KAIT data (purple) and *Swift* data (green) for SN 2009ig are displayed in different colors. These data from Foley et al. (2012a) have been corrected for systematic differences in their respective instrumental response functions using spectrophotometry (also known as an S -correction; e.g Stritzinger et al. 2002) that makes them slightly bluer than without the corrections. In order to avoid introducing additional systematics, we do not apply S -corrections to any of the photometry of SN 2012cg. The effect would be ≈ 0.1 mag at -16 days and the applied S -correction would make the early colors of SN 2012cg even bluer. No significant difference is apparent between the *Swift* colors and colors from ground based sources.

In contrast, the $B - V$ colors for the interaction models start very blue in the earliest phases. They grow rapidly redder and then approach the color curve of non-interacting SN Ia after a day or two. The $B - V$ color curves for the models reach a red peak within a few days and then they decline in unison with the color curve for a non-interacting SN Ia.

The early $B - V$ colors of SN 2012cg behave like the interaction models. The data have more scatter than the models, and they might show a peak near -15 days, but given the uncertainties they also look consistent with a flat color evolution up to -12 days after which they start to decline. In between -18 and -15 days the $B - V$ data of SN 2012cg are found near the model results for interactions a $6 M_{\odot}$ MS star. Irrespective of the best model fit, it is clear that SN 2012cg is bluer than non-interacting SNe Ia, either model or observed, but after -14 days the $B - V$ colors follow the path of a normal SN Ia.

These early photometric data of SN 2012cg are consistent with the interaction models and are clearly unlike SN 2009ig, SN 2011fe, or the model for a normal SN Ia that experiences no interaction. The model for a larger companion, such as the Red Giant, does not fit the timing of the data peak.

5. SPECTRA

The spectral features of SN 2012cg are similar to other SN Ia that are slightly overluminous with moderate decline rates. All of the typical SN Ia features are present, but in the earliest spectra, the velocities are slightly higher and features

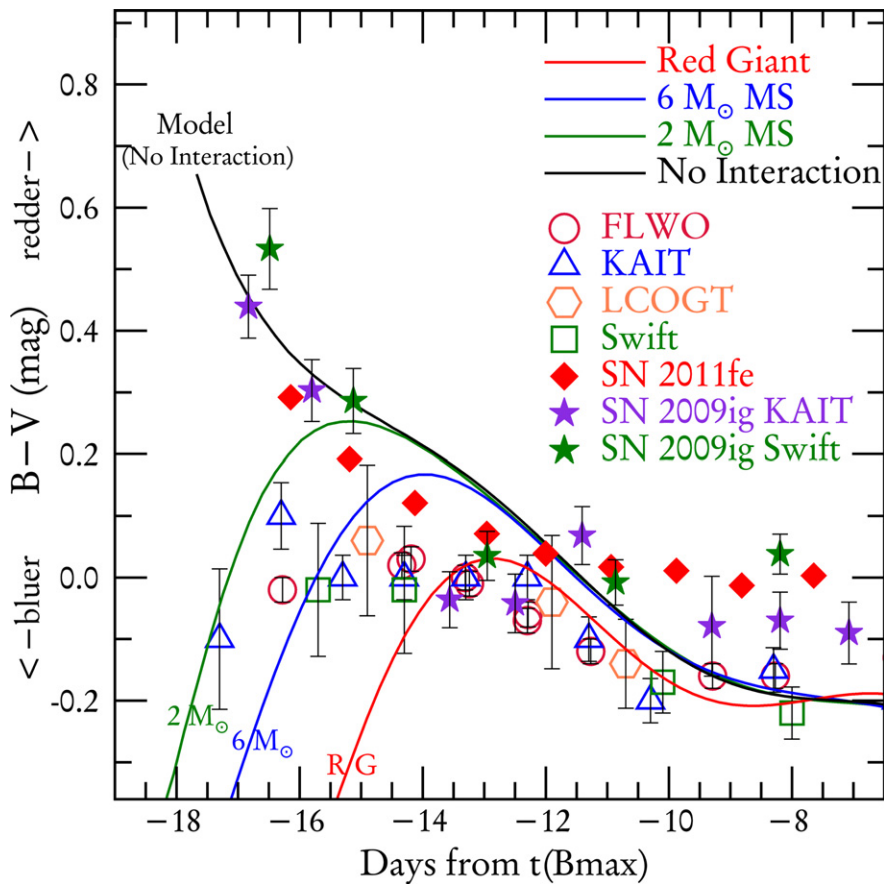


Figure 7. $B - V$ colors for the SN 2012cg are displayed using open symbols. The colored lines trace $B - V$ for the interaction models. $B - V$ data for two normal SN Ia are plotted using filled symbols: SNe 2009ig (Foley 2013, stars) and 2011fe (Pereira et al. 2013, diamonds). The colors for the other two SN are similar to the model for a normal SN Ia with no interaction (solid black line), but they do not fit the data for SN 2012cg or any of the interaction models. The colors for SN 2012cg are much bluer at the earliest phases, they redden for 2–3 days and then “turn over” near –15 days. After reaching this peak, the SN 2012cg colors follow the normal track. The $B - V$ data for SN 2012cg are closest to the model with a $6 M_{\odot}$ companion (blue line).

are slightly shallower than features found in spectra from fainter SN Ia. By –7 days, the spectra of SN 2012cg are similar to all normal SN Ia.

As shown by Marion et al. (2013) and Marion et al. (2015), plotting the pre-maximum features of multiple ions in the same velocity space is a productive tool to identify the relative locations of their line-forming regions in radial space. In particular, this technique allows easy identification and comparison of photospheric-velocity features (PVFs)—absorption features with minima indicating typical SN Ia photospheric velocities—and detached, HVFs—absorption features with minima indicating significantly higher velocities than typical SN Ia photospheric velocities (e.g., Marion et al. 2013; Childress et al. 2014; Silverman et al. 2015).

Figure 8 shows the pre-maximum spectra from Figure 3 zoomed in on individual absorption features of Ca II H&K, Si II $\lambda 6355$, C II $\lambda 6580$ and the Ca II infrared triplet (IR3; $\lambda 8579$). The phases of observation are from –16 days (top) to +0 days (bottom). The features have been normalized to a flat continuum, and the line depths have been normalized to 1.0 in arbitrary flux units.

The only measurable C II detection is for $\lambda 6580$ at –16 days. That relatively narrow absorption feature is plotted in black and can also be seen in the red wing of Si II $\lambda 6355$ at this phase (near 5000 km s^{-1}). The C II velocity of about $18,000 \text{ km s}^{-1}$ is clearly lower than the HVF for Si II. Silverman et al. (2012)

trace C II in SN 2012cg through about –8 days. They use SYN++ fitting (Thomas et al. 2011)²³ to tease measurements out of small distortions in the spectra.

Si II and Ca II show HVF components that change with time. Si II $\lambda 6355$ has a distinct HVF at –16 days with a velocity of about $-21,500 \text{ km s}^{-1}$ at the absorption minimum. From –15 to –12 days, the Si II HVF and PVF are blended into a single broad absorption feature. By –11 days, there is no longer evidence for Si II HVF. Si II PVF are measured to be $-10,500 \text{ km s}^{-1}$ on –9 days and the Si II velocities remain constant through $t(B_{\max})$.

HVF for Ca II are present in the spectra of most SN Ia. Here, they are prominent from –16 days to about –11 days. HVF Ca II becomes weaker compared to Ca II PVF, but they persist through $t(B_{\max})$. Near –10 days, the Ca II HVF feature becomes narrower, the limit of the blue wing moves from about $30,000 \text{ km s}^{-1}$ to about $25,000 \text{ km s}^{-1}$ and it makes a more abrupt transition to the continuum.

Ca II PVF are first detected in H&K beginning about –9 days. This phase is also when the primary absorption in Ca II H&K (seen in the figure at $-21,000 \text{ km s}^{-1}$) begins to be distorted by Si II $\lambda 3858$ (Foley 2013; Marion et al. 2013). The absorption minimum for these features is about $-11,000 \text{ km s}^{-1}$ in the rest frame of the Si II line, which suggests that it is indeed PVF Si II and not HVF Ca II.

²³ <https://c3.lbl.gov/es/>

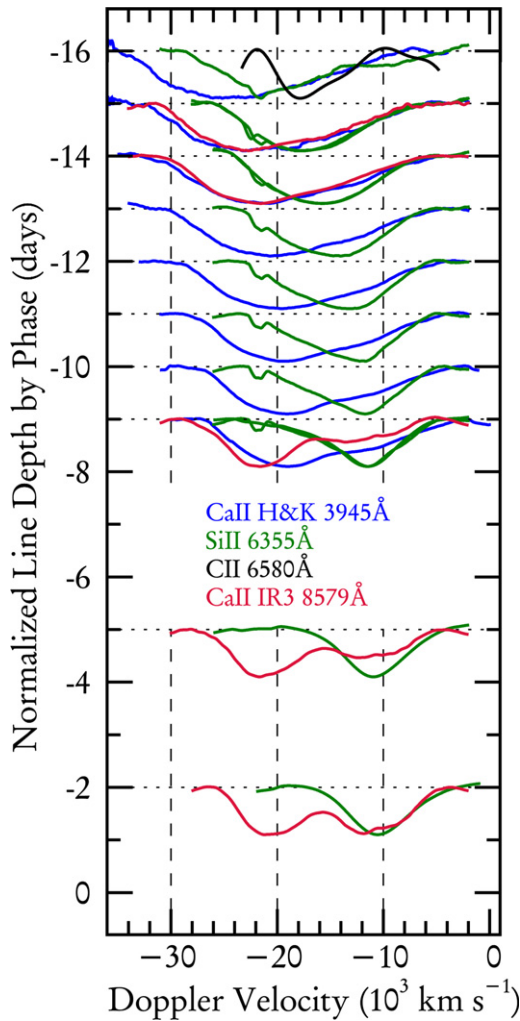


Figure 8. The evolution of absorption features C II (black), Si II (green), and Ca II (blue and red) in pre-maximum spectra of SN 2012cg. The features are normalized to a flat continuum and the line depths are normalized to 1.0. These line profiles identify the locations of line-forming regions for each ion in velocity space. C II is only measurable at -16 days. Si II exhibits a detached HVF at -16 days and strong influence of the HVF in Si II line profiles at -14 and -13 days. HVF of Ca II are common in SN Ia, and here they persist until at least -2 days.

5.1. Early Spectra are Very Blue

Figure 9 shows the FLWO optical spectra of SN 2012cg obtained at -16.1 , -15.0 and -14.2 days. The spectra are dereddened and plotted in blue when using the combined extinction of the host and MW ($E(B - V) = 0.20$), and plotted in green for MW extinction only ($E(B - V) = 0.018$). Reddening details can be found in Section 2.

The SN 2012cg spectra are compared to spectra of SN 2011fe that represent a typical SN Ia. They were obtained at -16.1 and -15.3 days (Parrent et al. 2012) and at -14.3 days (Pereira et al. 2013). After correction for extinction, the fluxes are scaled to the distance of SN 2012cg and plotted in red.

The spectra plotted in black are produced by adding the flux of a blackbody (BB) to the spectra from SN 2011fe. A different BB is added to the spectra from SN 2011fe to form the orange spectra. The BB parameters are listed in each panel. The SN 2012cg spectra are bluer and hotter than spectra of

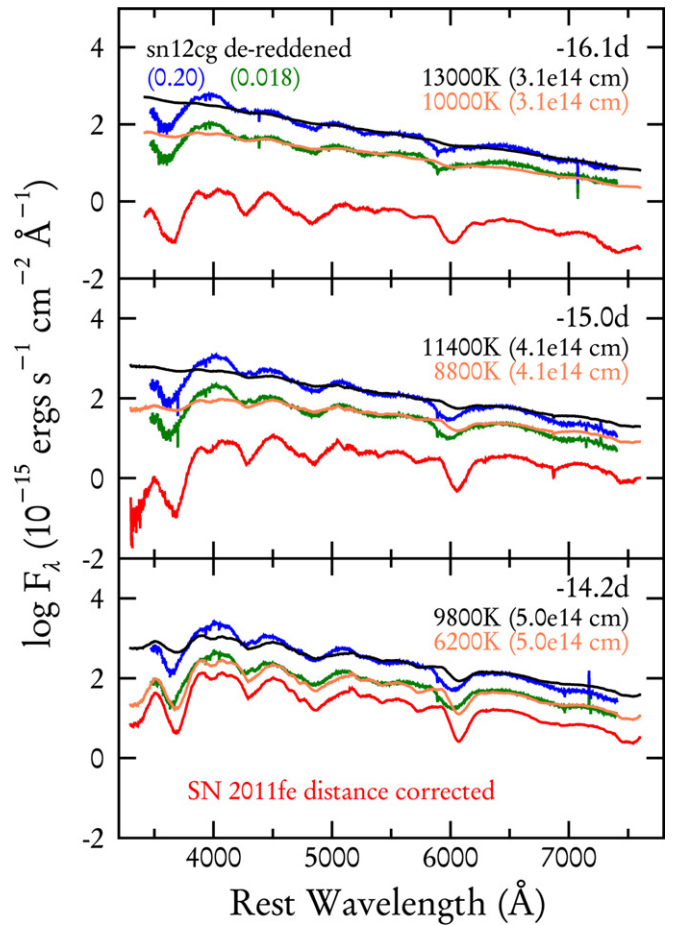


Figure 9. The earliest spectra of SN 2012cg are dereddened and plotted in blue for Host + MW extinction and green for MW only. The spectra plotted in red are distance corrected SN 2011fe spectra from similar phases. The black and orange spectra are the SN 2011fe data with blackbodies added to the temperatures and radii listed in each panel. Dilution of the features consistent with adding continuum to the spectra of SN 2011fe. Note that the BB radii are approximately the same as the photospheric radii at these phases. The SN 2012cg spectra are bluer and hotter than spectra of SN 2011fe at comparable phases which is consistent with the photometric measurements. At -14.2 days, the red spectrum is offset by -0.2 log flux units to make it easier to see the orange and green spectra.

SN 2011fe at comparable phases. This result is consistent with the photometric measurements.

Fitting BB shapes to the spectra does not generate precise physical measurements (Kirshner et al. 1973). The BB parameters displayed here are chosen so that when they are added to the SN 2011fe spectra, the resulting continua fit the spectra of SN 2012cg. The fitting was done by eye, but changes to the temperature or radius of less than 10% are sufficient to make it obvious that the results do not fit the target spectrum.

The radii for the BB that produce the closest fits to the SN 2012cg data are similar to the radii for the photosphere in homologous expansion. For example, we can define a photospheric radius using the time interval from the estimated time of explosion (-18.8 days) to the time of a observation. The rates of expansion, estimated from the absorption minima of the Si II 6355 feature (see Figure 8), are $\sim 16,000$ km s $^{-1}$ at -15.0 days and $\sim 14,000$ km s $^{-1}$ at -14.2 d. Since the strong HVF component of the Si II 6355 feature prevents reliable measurement of the photospheric component in the -16.1 days

spectrum, we use $\sim 16,000 \text{ km s}^{-1}$ from the spectrum taken 1 day later (-15.0 days).

The expansion parameters that determine the photospheric radius corresponding to the earliest spectrum would be: $t_1 = 2.7d = 2.33 \times 10^5 \text{ s}$ and $v_1 = 16 \times 10^9 \text{ cm s}^{-1}$. Therefore at a phase of -16.1 days, we estimate the radius of the photosphere to be: $R_1 \approx 3.7 \times 10^{14} \text{ cm}$. For the phase of -15.0 days, the velocity stays the same and $R_2 \approx 5.2 \times 10^{14} \text{ cm}$. At -14.4 days, we apply $v_3 = 14 \times 10^9 \text{ cm s}^{-1}$ to the time interval from -15.0 days, and find that $R_3 \approx 6.1 \times 10^{14} \text{ cm}$.

High temperature blackbodies, as in the top panel, wash out the spectral features of the SN 2011fe spectra. Lower temperatures, as found in the bottom panel, make only small changes to the features. The stronger features in SN 2012cg imply that the excess flux is not a pure BB. The SN 2012cg features appear to have a higher optical depth at these very early phases than they do at later phases. This means higher density ejecta in the line-forming regions at -16 and -15 days. By -14.4 days the spectral differences are small.

5.2. Narrow Features from the Host

Significant changes to narrow features from Na I D and Ca II H&K may be evidence of a process that could produce an observable increase in luminosity. If the SN shock passes through a region of CSM that has a higher density than is usually found around SN Ia, then Na I and Ca II would be ionized. As the region cools, the atoms recombine and may produce additional luminosity. The recombination will also increase the absorption strength of the observed Na I and Ca II lines (Patat et al. 2007; Blondin et al. 2009, 2012; Simon et al. 2009).

These narrow features are strong in all spectra in this sample, but the resolution of the FAST spectra ($R \approx 2700$) is not intended for measurements on the order of a few hundred kilometers per second. The pseudo equivalent widths of the Ca II and Na I lines fluctuate within the expected uncertainties due to random noise. There is no evidence that the line profiles or equivalent widths are any different during the phases at which we measure excess flux (-16 and -15 days) than they are at any other phase in our sample. Thus, we rule out the possibility that the observed luminosity enhancement is due to interaction with circumstellar material.

6. DISCUSSION AND CONCLUSIONS

We find excess luminosity at -16 and -15 days in the LCs of SN Ia 2012cg. The excess is present in data from multiple filters obtained at multiple sources. The $B - V$ color curves for SN 2012cg are very blue at these phases, and they clearly diverge from $B - V$ models and data of normal SN Ia. Spectral evidence is also used to confirm that this short period of excess luminosity, just a few days after the explosion, is real.

The basic parameters for SN 2012cg describe a slightly overluminous SN Ia with a moderately slow decline rate and a normal Si II velocity: $M_B = -19.62 \text{ mag}$, $\Delta m_{15}(B) = 0.86 \text{ mag}$ and $v_{\text{Si}} = -10, 500 \text{ km s}^{-1}$ at $t(B_{\text{max}})$. These parameters agree with Silverman et al. (2012) who were unable to detect the extra flux because their earliest observations were near the detection limit.

We also examine optical and infrared spectra from SN 2012cg that reveal a blue continuum and relatively weak

absorption features in the pre-maximum spectra. HVF are detected for both Si II and Ca II. Silicon velocities are normal ($v_{\text{Si}} = -10, 500 \text{ km s}^{-1}$) at -10 days and they stay constant through $t(B_{\text{max}})$.

Kasen (2010) described how the impact of the SN on a companion could produce emission that may be detected as enhanced luminosity in the first few days after the explosion. We compare the early LC data and $B - V$ colors of SN 2012cg to the Kasen (2010) models for a normal SN Ia and to models for interaction with three different, non-degenerate companions. We find that the observations of SN 2012cg are consistent with models for the interaction between a SN Ia and a main-sequence binary companion of about $6 M_{\odot}$.

Interaction with a larger companion star is a possibility with the constraints of our data, if the impact site were significantly off-axis with respect to the line of sight to the SN. The interaction would still produce excess flux but at a reduced level that may be comparable to the observations (Kasen 2010; Brown et al. 2012).

Our size estimate for the companion is supported by Graur et al. (2015). They use pre-explosion *HST* Wide-Field Planetary Camera 2 images to estimate the upper limits on the luminosity of a possible companion to SN 2012cg. The limits they derive suggest that the brightest possible companion would be either a Red Giant or a $\approx 7 M_{\odot}$ main-sequence star. The late-time *HST* WFC3 photometry from Graur et al. (2015) is also consistent with this model. The pre-explosion limits suggest that the progenitor system of SN 2012cg did not have a helium star donor (e.g., Wang et al. 2014; Liu et al. 2010).

We note that Dessart et al. (2014) present pulsational-delayed detonation (PDD) models of SN Ia, and some of these models exhibit color peaks similar to the Kasen (2010) models. The PDD models also have increased luminosity at early times, but they do not reach the levels of excess that we observe in SN 2012cg. Comparing the $B - V$ colors of SN 2012cg, SN 2011fe, and SN 2009ig, as plotted in Figure 7, to the results from synthetic photometry of various PDD models²⁴ by Dessart et al. (2014) we find that the PDD model colors are consistent with the observations at -14 days and afterward, while before this epoch the PDD models predict $B - V > 0.2 \text{ mag}$, similar to SN 2011fe and SN 2009ig but in disagreement with the bluer ($B - V \sim 0 \text{ mag}$) colors of SN 2012cg.

Other models predict that extra blue flux at early times could be the result of radioactive ^{56}Ni on the outside of the exploding WD. This material could potentially come from the burning of accreted He on the surface WD progenitor. Shen & Moore (2014) discuss He detonation models for single and DD progenitors of SN Ia. They find that the He detonation would likely be triggered by a minimum mass He shell. In the case of a low mass shell, He will only burn to Si and Ca, and not produce radioactive Ni. Although a larger He shell may burn all the way to ^{56}Ni , Shen & Moore (2014) predict that the He detonation will occur at lower masses and trigger a SN Ia well before the He-shell becomes capable of producing ^{56}Ni .

Amanullah et al. (2015) show that SN 2012cg may have time variable reddening that would change the colors with time. However, this effect is reported in only the $M2 - V$ colors and only nearer to $t(B_{\text{max}})$ than the phases we discuss here. They note that $B - V$ has minimal evolution and they specifically

²⁴ https://www-n.oca.eu/supernova/sn1a/sn1a_ddc_pddel.html

discount the presence of circumstellar dust close enough to the SN that it might affect the brightness and colors we observe at these very early phases. Thus, the unusual early-time colors we see are probably not due to CSM interaction, but instead are coming from interaction with the companion.

Cortes et al. (2006) report that NGC 4424 has a strongly disturbed stellar disk, with groups of young blue stars outside the locations of current star formation. They suggest that the peculiarities of NGC 4424 are the result of an intermediate-mass merger plus ram pressure stripping. SN 2012cg went off on the East side of the host, about $17''$ away from the nucleus. That puts it outside the current H α emission region, but within regions that mix dust, blue-star-complexes and H I in emission. SN 2012cg is not in a region of active, ongoing star formation, but it is in a location where there was star formation in the recent past.

Crowl & Kenney (2008) describe fiber spectroscopy of NGC 4424. Their fibers clearly encompass the location of SN 2012cg. All of the fiber spectra are averaged to obtain a composite spectrum of the galaxy. They determine that the luminosity weighted stellar population of the composite spectrum is about 2 Gyr and that star formation shut off about 300-500 Myr ago.

The turbulent region around SN 2012cg makes it difficult to be precise about the star formation history. Given these constraints, the maximum mass of stars currently present in this location is likely to approach ZAMS $\approx 9M_{\odot}$. They will have been formed toward the end of the recent star formation epoch. Less massive stars are also possible. Thus, the environment is suitable to establish a progenitor system in the region of SN 2012cg with a WD of about $1M_{\odot}$ and a MS companion of about $6M_{\odot}$.

This work emphasizes the importance of photometric observations of SN Ia as early as -17 days to evaluate the possible interaction with a companion and for more advanced analysis of progenitor systems. Foley et al. (2012a) identified excess UV flux from SN 2009ig at early phases, but the colors were inconsistent with an interaction. SN 2011fe was observed early but the data reveal no evidence for interaction. Nearly all of the LC of other SN Ia that have been used to interpret the presence or absence of interaction do not include sufficiently early data. However, these reported non-detections have been used often to suggest that the SD model for SN Ia is no longer viable and that SN Ia are exclusively produced in DD progenitor systems.

Observations at phases between -17 and -15 days suggest that SN 2012cg had a MS binary companion of about $6M_{\odot}$ when it exploded. This is consistent with the SD scenario, i.e., a binary system containing the exploding white dwarf and a non-degenerate companion. The size estimate for the companion should be regarded as a minimum, since the star could be much larger but seen at a less favorable angle. The constraints of timing and the size of the companion demonstrate that the interaction with a companion would go undetected in all but a few of the current data sets. Had the angle of observation been different, the excess luminosity in SN 2012cg might have gone unobserved as well.

J.V. is supported by Hungarian OTKA grant NN-107637. J.C.W., G.H.M., J.M.S. and the U.T. supernova group are supported by NSF grant AST-1109801. J.M.S. is also supported by an NSF Astronomy and Astrophysics

Postdoctoral Fellowship under award AST-1302771. The CfA Supernova Program is supported by NSF grants AST-1211196 and AST-156854 to the Harvard College Observatory. R.P.K. was supported in part by the National Science Foundation under Grant NSF PHY-1125915 to the Kavli Institute for Theoretical Physics. A.S.F. acknowledges support from a NSF Graduate Research Fellowship, a NASA Graduate Research Program Fellowship, and a NSF STS postdoctoral fellowship under award SES-1056580. We acknowledge the work of C. Klein, D. L. Starr, and J. S. Bloom on the PAIRITEL mosaic data reduction pipeline. Additional support comes from program GO-12540, provided by NASA through a grant from the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS5-26555. ROTSE data analysis was supported by NASA grant NNX10A196H subaward UTA13-000844.

G.H.M. is a visiting Astronomer at the Infrared Telescope Facility, which is operated by the University of Hawaii under Cooperative Agreement no. NNX-08AE38A with the National Aeronautics and Space Administration, Science Mission Directorate, Planetary Astronomy Program. P. J. Brown and the *Swift* Optical/Ultraviolet Supernova Archive are supported by NASA's Astrophysics Data Analysis Program through grant NNX13AF35G. E.Y.H. acknowledges the support provided by the National Science Foundation under grant no. AST-1008343 and by the Danish Agency for Science, Technology and Innovation through a Sapere Aude Level 2 grant. This research at Rutgers University was supported by NSF CAREER award AST-0847157 to S.W.J., and NSF REU grant PHY-1263280 for Y.C. The work by K. Maeda is supported by JSPS KAKENHI (26800100) and MEXT WPI Initiative, Japan. This paper includes data gathered with the 6.5 m *Magellan* Telescopes located at the Las Campanas Observatory, Chile. This work makes use of observations from the LCOGT network. The authors make frequent use of David Bishop's excellent webpage listing recent supernovae and valuable references associated with them: www.rochesterastromy.org/snimages/.

Facilities: FLWO:1.5m (FAST), HET (LRS), Swift (UVOT; UV grism), IRTF (SpeX), FLWO:1.2m (KepCam), LCOGT, Lick (KAST), SALT (RSS).

REFERENCES

- Amanullah, R., Johansson, J., Goobar, A., et al. 2015, *MNRAS*, 453, 3300
 Arnett, W. D. 1982, *ApJ*, 253, 785
 Bianco, F. B., Howell, D. A., Sullivan, M., et al. 2011, *ApJ*, 741, 20
 Blondin, S., Matheson, T., Kirshner, R. P., et al. 2012, *AJ*, 143, 126
 Blondin, S., Prieto, J. L., Patat, F., et al. 2009, *ApJ*, 693, 207
 Breeveld, A. A., Landsman, W., Holland, S. T., et al. 2011, in AIP Conf. Ser. 1358, Gamma Ray Bursts 2010, ed. J. E. McEnery, J. L. Rascusin, & N. Gehrels (Melville, NY: AIP), 373
 Brown, P. J., Breeveld, A. A., Holland, S., et al. 2014, *Ap&SS*, 354, 89
 Brown, P. J., Dawson, K. S., & de Pasquale, M. 2012, *ApJ*, 753, 22
 Brown, P. J., Holland, S. T., Immler, S., et al. 2009, *AJ*, 137, 4517
 Brown, T. M., Baliber, N., Bianco, F. B., et al. 2013, *PASP*, 125, 1031
 Cao, Y., Kulkarni, S. R., Howell, D. A., et al. 2015, *Natur*, 521, 328
 Cenko, S. B., Filippenko, A. V., Silverman, J. M., et al. 2012, *ATel*, 4115, 1
 Childress, M. J., Filippenko, A. V., Ganeshalingam, M., et al. 2014, *MNRAS*, 437, 338
 Chomiuk, L. 2013, *PASA*, 30, 46
 Chomiuk, L., Soderberg, A. M., Moe, M., et al. 2012, *ApJ*, 750, 164
 Chonis, T. S., & Gaskell, C. M. 2008, *AJ*, 135, 264
 Cortes, J. R., Kenney, J. D. P., & Hardy, E. 2006, *AJ*, 131, 747
 Cortes, J. R., Kenney, J. D. P., & Hardy, E. 2008, *ApJ*, 683, 78
 Crawford, S. M., Still, M., Schellart, P., et al. 2010, *Proc. SPIE*, 7737E, 25

- Crowl, H. H., & Kenney, J. D. P. 2008, *AJ*, 136, 1623
- Cutri, R. M., Skrutskie, M. F., van Dyk, S., et al. 2003, tmc, book, The IRSA 2MASS All-Sky Point Source Catalog NASA/IPAC Infrared Science Archive, <http://irsa.ipac.caltech.edu/applications/Gator>
- Dessart, L., Blondin, S., Hillier, D. J., et al. 2014, *MNRAS*, 441, 532
- Dilday, B., Howell, D. A., Cenko, S. B., et al. 2012, *Sci*, 337, 942
- Fabricant, D., Cheimets, P., Caldwell, N., & Geary, J. 1998, *PASP*, 110, 79
- Filippenko, A. V., Li, W. D., Treffers, R. R., & Modjaz, M. 2001, in ASP Conf. Ser. 246, Small Telescope Astronomy on Global, ed. B. Paczynski, W. P. Chen, & C. Lemme (San Francisco, CA: ASP), 121
- Fitzpatrick, E. L. 1999, *PASP*, 111, 63
- Foley, R. J. 2013, *MNRAS*, 435, 273
- Foley, R. J., Challis, P. J., Filippenko, A. V., et al. 2012a, *ApJ*, 744, 38
- Foley, R. J., Simon, J. D., Burns, C. R., et al. 2012b, *ApJ*, 752, 101
- Friedman, A. S. 2012, PhD thesis, Harvard Univ.
- Friedman, A. S., Wood-Vasey, W. M., Marion, G. H., et al. 2015, *ApJS*, 220, 9
- Ganeshalingam, M., Li, W., & Filippenko, A. V. 2011, *MNRAS*, 416, 2607
- Ganeshalingam, M., Li, W., Filippenko, A. V., et al. 2010, *ApJS*, 190, 418
- Ganeshalingam, M., Li, W., Filippenko, A. V., et al. 2012, *ApJ*, 751, 142
- Goobar, A., Kromer, M., Siverd, R., et al. 2015, *ApJ*, 799, 106
- Graur, O., Maoz, D., & Shara, M. M. 2014, *MNRAS*, 442L, 28
- Graur, O., Zurek, D., Shara, M. M., et al. 2015, in press
- Hayden, B. T., Garnavich, P. M., Kasen, D., et al. 2010, *ApJ*, 722, 1691
- Hicken, M., Challis, P., Kirshner, R. P., et al. 2012, *ApJS*, 200, 12
- Hicken, M., Garnavich, P. M., Prieto, J. L., et al. 2007, *ApJL*, 669, L17
- Hill, G. J., Nicklas, H. E., MacQueen, P. J., et al. 1998, Proc. SPIE, 3355, 375
- Hillebrandt, W., & Niemeyer, J. C. 2000, *ARA&A*, 38, 191
- Iben, I., & Tutukov, A. V. 1984, *ApJS*, 54, 335
- Kasen, D. 2010, *ApJ*, 708, 1025
- Kashi, A., & Soker, N. 2011, *MNRAS*, 417, 1466
- Kirshner, R. P., Willner, S. P., Becklin, E. E., et al. 1973, *ApJ*, 180L, 97
- Kushnir, D., Katz, B., Dong, S., et al. 2013, *ApJ*, 778L, 37
- Kutsuna, M., & Shigeyama, T. 2015, *PASJ*, 67, 54
- Landolt, A. U. 1992, *AJ*, 104, 340
- Li, W., Bloom, J. S., Podsiadlowski, P., et al. 2011, *Natur*, 480, 348
- Liu, W.-M., Chen, W.-C., Wang, B., & Han, Z. W. 2010, *A&A*, 523A, 3
- Maeda, K., Kutsuna, M., & Shigeyama, T. 2014, *ApJ*, 794, 37
- Maguire, K., Sullivan, M., Patat, F., et al. 2013, *MNRAS*, 436, 222
- Mandel, K., Narayan, G., & Kirshner, R. P. 2011, *ApJ*, 731, 120
- Maoz, D., Mannucci, F., & Nelemans, G. 2014, *ARA&A*, 52, 107
- Margutti, R., Parrent, J., Kamble, A., et al. 2014, *ApJ*, 790, 52
- Margutti, R., Soderberg, A. M., Chomiuk, L., et al. 2012, *ApJ*, 751, 134
- Marietta, E., Burrows, A., & Fryxell, B. 2000, *ApJS*, 128, 615
- Marion, G. H., Sand, D. J., Hsiao, E. Y., et al. 2015, *ApJ*, 798, 39
- Marion, G. H., Vinkó, J., Wheeler, J. C., et al. 2013, *ApJ*, 777, 40
- Matheson, T., Blondin, S., Foley, R. J., et al. 2005, *AJ*, 129, 2352
- Miknaitis, G., Pignata, G., Rest, A., et al. 2007, *ApJ*, 666, 674
- Monet, D. G., Levine, S. E., Canzian, B., et al. 2003, *AJ*, 125, 984
- Munari, U., Henden, A., Belligoli, R., et al. 2013, *NewA*, 20, 30
- Nomoto, K. 1982, *ApJ*, 253, 798
- Nugent, P. E., Sullivan, M., Cenko, S., et al. 2011, *Natur*, 480, 344
- Olling, R. P., Mushotzky, R., Shaya, E. J., et al. 2015, *Natur*, 521, 332
- Parrent, J. T., Howell, D. A., Friesen, B., et al. 2012, *ApJ*, 752L, 26
- Patat, F., Chandra, P., Chevalier, R., et al. 2007, *Sci*, 317, 924
- Pereira, R., Thomas, R. C., Aldering, G., et al. 2013, *A&A*, 554, A27
- Phillips, M. M. 1993, *ApJ*, 413L, 105
- Piro, A. L., & Nakar, E. 2013, *ApJ*, 769, 67
- Ramsey, L. W., Adams, M. T., Barnes, T. G., et al. 1998, Proc. SPIE, 3352, 34
- Roming, P. W. A., Kennedy, T. E., & Mason, K. O. 2005, *SSRv*, 120, 95
- Scalzo, R., Aldering, G., Antilogus, P., et al. 2014, *MNRAS*, 440, 1498
- Schaefer, B. E., & Pagnotta, A. 2012, *Natur*, 481, 164
- Schechter, P. L., Mateo, M., & Saha, A. 1993, *PASP*, 105, 1342
- Schlafly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, 737, 103
- Shen, K., & Moore, K. 2014, *ApJ*, 797, 46
- Silverman, J. M., Ganeshalingam, M., Cenko, S. B., et al. 2012, *ApJL*, 756, 7
- Silverman, J. M., Vinkó, J., Marion, G. H., et al. 2015, *MNRAS*, 451, 1973
- Simon, J. D., Gal-Yam, A., Gnat, O., et al. 2009, *ApJ*, 702, 1157
- Smith, J. A., Tucker, D. L., Kent, S., et al. 2002, *AJ*, 123, 2121
- Sternberg, A., Gal-Yam, A., Simon, J. D., et al. 2011, *Sci*, 333, 856
- Stritzinger, M., Hamuy, M., Suntzeff, N. B., et al. 2002, *AJ*, 124, 2100
- Thomas, R. C., Nugent, P. E., & Meza, J. C. 2011, *PASP*, 123, 237
- Tucker, B. E. 2011, *Ap&SS*, 335, 223
- Vinkó, J., Sárneczky, K., Takáts, K., et al. 2012, *A&A*, 546, 12
- Wang, B., Meng, X., Liu, D.-D., et al. 2014, *ApJ*, 794L, 28
- Webbink, R. F. 1984, *ApJ*, 277, 355
- Whelan, J., & Iben, I. 1973, *ApJ*, 186, 1007
- Yaron, O., & Gal-Yam, A. 2012, *PASP*, 124, 668
- Yuan, F., & Akerlof, C. W. 2008, *ApJ*, 667, 808
- Zheng, W., Shivvers, I., Filippenko, A. V., et al. 2014, *ApJ*, 783L, 24
- Zheng, W., Silverman, J. M., Filippenko, A. V., et al. 2013, *ApJ*, 778, 15