

SPTpol: an instrument for CMB polarization measurements with the South Pole Telescope

J.E. Austermann^{2,3}, K.A. Aird²⁰, J.A. Beall⁴, D. Becker⁴, A. Bender⁶, B.A. Benson^{1,21}, L.E. Bleem^{1,19}, J. Britton⁴, J.E. Carlstrom^{1,5,19,21,22}, C.L. Chang^{1,5,21}, H.C. Chiang^{1,21}, H-M. Cho⁴, T.M. Crawford^{1,22}, A.T. Crites^{1,22}, A. Datesman⁷, T. de Haan⁶, M.A. Dobbs⁶, E.M. George⁸, N.W. Halverson^{2,3,18}, N. Harrington⁸, J.W. Henning^{2,3}, G.C. Hilton⁴, G.P. Holder⁶, W.L. Holzappel⁸, S. Hoover^{1,19}, N. Huang^{1,21}, J. Hubmayr⁴, K.D. Irwin⁴, R. Keisler^{1,19,21}, J. Kennedy⁶, L. Knox²³, A.T. Lee⁸, E. Leitch¹, D. Li⁴, M. Lueker¹⁰, D.P. Marrone¹⁷, J.J. McMahon¹¹, J. Mehl^{1,21}, S.S. Meyer^{1,19,21,22}, T.E. Montroy¹³, T. Natoli^{1,19}, J.P. Nibarger⁴, M.D. Niemack⁴, V. Novosad⁷, S. Padin¹, C. Pryke¹², C.L. Reichardt⁸, J.E. Ruhl¹³, B.R. Saliwanchik¹³, J.T. Sayre¹³, K.K. Schaffer¹⁴, E. Shirokoff¹⁰, A.A. Stark²⁴, K. Story^{1,19}, K. Vanderlinde⁶, J.D. Vieira¹⁰, G. Wang⁵, R. Williamson^{1,21}, V. Yefremenko^{5,7}, K. W. Yoon⁴, O. Zahn⁸

¹ Kavli Institute for Cosmological Physics, Department of Physics, Enrico Fermi Institute, The University of Chicago, Chicago, IL 60637, USA

² Department of Astrophysical and Planetary Sciences, University of Colorado, Boulder, Colorado, 80309, USA

³ Center for Astrophysics and Space Astronomy, University of Colorado, Boulder, Colorado, 80309, USA

⁴ NIST, Boulder, CO 80305, USA

⁵ High Energy Physics Division, Argonne National Laboratory, Argonne, IL 60439, USA

⁶ McGill University, Montreal, Quebec, Canada

⁷ Materials Science Division, Argonne National Laboratory, Argonne, IL 60439, USA

⁸ University of California, Berkeley, 151 LeConte Hall Berkeley, CA 94720, USA

¹⁰ California Institute of Technology, Pasadena, CA 91125, USA

¹¹ University of Michigan, Ann Arbor, Michigan, USA

¹² University of Minnesota, Minneapolis, MN 55455, USA

¹³ Case Western Reserve University, Cleveland, Ohio 44106, USA

¹⁴ School of the Art Institute of Chicago, Chicago, Illinois, 60603, USA

¹⁷ Steward Observatory, University of Arizona, 933 North Cherry Avenue, Tucson, AZ 85721, USA

¹⁸ Department of Physics, University of Colorado, Boulder, CO 80309

¹⁹ Department of Physics, University of Chicago, 5640 South Ellis Avenue, Chicago, IL, USA 60637

²⁰ University of Chicago, 5640 South Ellis Avenue, Chicago, IL, USA 60637

²¹ Enrico Fermi Institute, University of Chicago, 5640 South Ellis Avenue, Chicago, IL, USA 60637

²² Department of Astronomy and Astrophysics, University of Chicago, 5640 South Ellis Avenue, Chicago, IL, USA 60637

²³ Department of Physics, University of California, One Shields Avenue, Davis, CA, USA 95616

²⁴ Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA, USA 02138

ABSTRACT

SPTpol is a dual-frequency polarization-sensitive camera that was deployed on the 10-meter South Pole Telescope in January 2012. SPTpol will measure the polarization anisotropy of the cosmic microwave background (CMB) on angular scales spanning an arcminute to several degrees. The polarization sensitivity of SPTpol will enable a detection of the CMB “B-mode” polarization from the detection of the gravitational lensing of the CMB by large scale structure, and a detection or improved upper limit on a primordial signal due to inflationary gravity waves. The two measurements can be used to constrain the sum of the neutrino masses and the energy scale of inflation. These science goals can be achieved through the polarization sensitivity of the SPTpol camera and careful control of systematics. The SPTpol camera consists of 768 pixels, each containing two transition-edge sensor (TES) bolometers coupled to orthogonal polarizations, and a total of 1536 bolometers. The pixels are sensitive to light in one of two frequency bands centered at 90 and 150 GHz, with 180 pixels at 90 GHz and 588 pixels at 150 GHz. The SPTpol design has several features designed to control polarization systematics, including: single-mode feedhorns with low cross-polarization, bolometer pairs well-matched to difference atmospheric signals, an improved ground shield design based on far-sidelobe measurements of the SPT, and a small beam to reduce temperature to polarization leakage. We present an overview of the SPTpol instrument design, project status, and science projections.

Keywords: Cosmology, TES, bolometer, millimeter, polarimetry, instrumentation

1. INTRODUCTION

The South Pole Telescope (SPT) is a 10-meter mm-wavelength telescope at the geographic South Pole. In November 2011, the SPT completed a 2500 deg² survey at 95, 150, and 220 GHz, the SPT-SZ survey, that has already led to significant results and new discoveries in three main areas: using the Sunyaev-Zel’dovich (SZ) effect to discover massive high-redshift galaxy clusters and constrain cosmological parameters,¹⁻⁵ measurements of fine-scale anisotropy of the cosmic microwave background (CMB) and the gravitational lensing of it from large scale structure,⁶⁻¹⁰ and the discovery of strongly lensed high-redshift star forming galaxies.¹¹

In January 2012, the SPT was equipped with a new polarization-sensitive camera, SPTpol. SPTpol is a dual-frequency polarization-sensitive camera sensitive to 90 and 150 GHz that will continue the original SPT science goals, and, in addition, measure the polarization anisotropy of the CMB on angular scales from an arcminute to several degrees. CMB polarization measurements are sensitive to multiple cosmological parameters that are related to open questions in fundamental physics, including the sum of the neutrino masses, the number of relativistic particle species at recombination, and the energy scale of inflation. The sensitivity of the SPTpol camera is expected to enable new constraints on each through its measurement of the power spectrum of the CMB polarization anisotropy.

This paper is organized as follows. In Section 2, we discuss the science goals for SPTpol. In Section 3, we describe the telescope and optics. In Section 4, we describe the design of the SPTpol camera, including the detectors and readout. In Section 5, we discuss the SPTpol survey strategy. Finally, in Section 6, we present the details of the SPTpol receiver performance, on-sky characterization, and the science projections for the SPTpol survey.

J.E.A.: E-mail: jason.austermann@colorado.edu

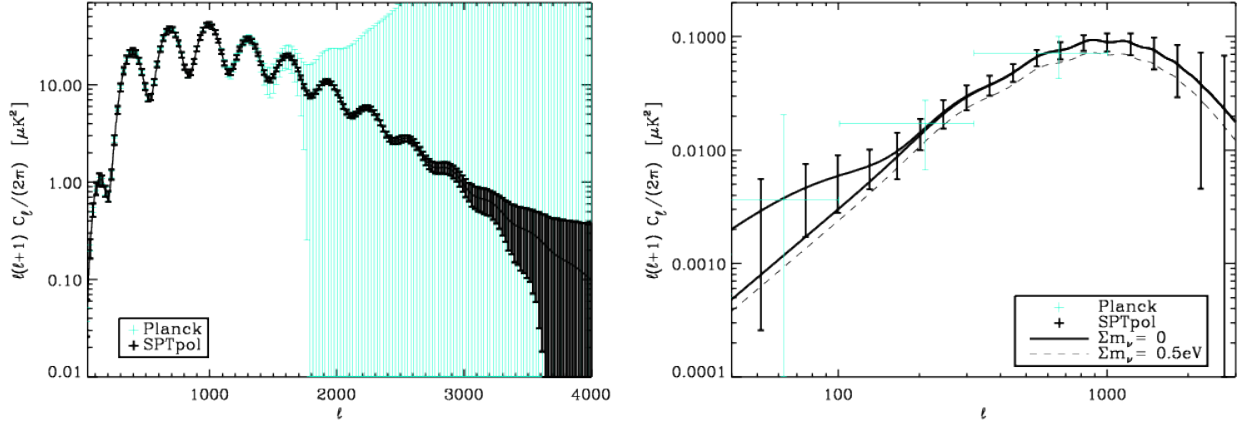


Figure 1. Projected EE (left) and BB (right) constraints from three years of observing with the SPTpol camera (black points and error bars). Constraints are from simulated observations including realistic treatment of foregrounds, atmosphere, instrument $1/f$ noise, and E - B separation while assuming detector performance listed in Table 3 with 80% pixel yield. Overplotted are projected constraints from *Planck*.¹² Solid model curves in the BB plot are for $\Sigma m_{\nu} = 0$, with $r = 0$ and $r = 0.04$, while the dashed model curve is for $\Sigma m_{\nu} = 0.5\text{eV}$ with $r = 0$.

2. SCIENCE GOALS

SPTpol will make sensitive measurements of the temperature and polarization anisotropy of the CMB on angular scales ranging from a few degrees to an arcminute, or angular multipoles between $\ell \sim 50 - 10,000$. The polarization signal can be decomposed into what are commonly referred to as “E-mode” (gradient-like) and “B-mode” (curl-like) signals; an allusion to the ‘E’ and ‘B’ field patterns in electromagnetism. These patterns are sourced and/or induced by distinct physical properties which can be separated through this decomposition. These measurements contain a wealth of information about the initial conditions, content, and evolution of the Universe, and can be used to constrain a host of fundamental cosmological parameters, including: the mass and number of neutrinos, the primordial power spectrum, the energy scale of inflation, and the density and equation of state of dark energy. In this section, we will discuss the science goals of the SPTpol experiment, and the projected cosmological constraints.

2.1 B-Mode Polarization Science

2.1.1 Inflationary B-modes

The only process expected to generate primordial B-mode polarization anisotropy is from gravity waves generated during the inflationary epoch, which are predicted to produce a signal at large angular scales ($\ell < 100$) with an amplitude related to the energy scale of inflation, and proportional to the tensor-to-scalar ratio, r , of the primordial perturbations.¹³ Detection of this signal is often considered both a confirmation of the theory of inflation, and tantalizing information about physics at energy scales far beyond those probed at terrestrial accelerators.¹⁴

While the detection of gravitational-wave B-modes will be extremely challenging, the raw sensitivity of SPTpol and the low-foreground sky continuously observable from the polar site will allow SPTpol to place a cosmologically compelling constraint on the gravitational-wave B-mode spectrum that could rule out various large- r inflationary models. The right panel of Figure 1 shows simulated B-mode spectra for the scenarios of $r = 0$ and $r = 0.04$. Using the simulated observations described below, we project that with three years of observations and realistic estimates of atmospheric contamination, the 1σ detection limit in the joint fit to gravity-wave and lensing template B-mode spectra is $r = 0.028$. This represents a significant improvement from the best current constraints on the tensor-to-scalar ratio from both the B-mode polarization, $r < 0.7$ at 95% confidence,¹⁵ and the CMB temperature anisotropy, $r < 0.21$ at 95% confidence.⁸

Dataset	Cosmological parameter constraints								
	$\sigma(\Omega_b h^2)$ $\times 10^4$	$\sigma(\Omega_c h^2)$ $\times 10^3$	$\sigma(A_s)$ $\times 10^{11}$	$\sigma(n_s)$ $\times 10^3$	$\sigma(h)$ $\times 10^2$	$\sigma(\tau)$ $\times 10^3$	$\sigma(N_{\text{eff}})$ $\times 10^1$	$\sigma(\Sigma m_\nu)$ [meV]	$\sigma(r)$ $\times 10^2$
<i>Planck</i>	1.93	2.02	5.36	7.07	1.88	4.96	1.39	117	5.72
+ SPTpol	1.64	1.71	4.92	6.19	1.58	4.95	1.17	96	2.75

Table 1. Expected 1σ constraints on cosmological parameters using SPTpol power spectrum and lensing reconstruction data, assuming a 9-parameter Λ CDM+ N_{eff} + Σm_ν +tensor model. In addition, when applied to a model including the primordial helium abundance (Λ CDM+ N_{eff} + Y_p + Σm_ν cosmology) SPTpol will improve the constraint on Y_{He} by a factor of ~ 2 over *Planck* alone.

2.1.2 Gravitational Lensing and Neutrino Masses

The other significant source of B-mode polarization is generated from the gravitational lensing of the CMB by large-scale structure, which converts some of the intrinsic E-mode polarization to B-modes. This contribution to the B-mode angular power spectrum is expected to peak around $\ell \sim 1000$ and dominate the primordial B-mode signal at $\ell > 200$ for any allowed value of the tensor-to-scalar ratio r , for example, see Figure 1. The shape and amplitude of the lensed B-mode power spectrum depends on the sum of neutrino masses, Σm_ν , because the energy density of massive neutrinos decreases more slowly with expansion than is the case for massless neutrinos. The resulting higher density, and thus higher universal expansion rate, leads to slower growth of structure — and, hence, suppression of lensed B-mode power — on scales smaller than the neutrino free-streaming length. SPTpol, with its high-resolution $1'$ beam, is capable of measuring the small-scale B-mode power spectrum with high precision, resulting in scientifically interesting constraints on Σm_ν .

In the right panel of Figure 1, we show that the SPTpol measurement of the B-mode power spectrum has the power to clearly distinguish between spectra with $\Sigma m_\nu = 0$ eV and $\Sigma m_\nu = 0.5$ eV. Including priors from *Planck*, see Table 1, we expect to constrain $\sigma(\Sigma m_\nu) = 0.096$ eV, a constraint that is $\sim 4\times$ better than future beta decay experiments such as KATRIN,¹⁶ which has a predicted sensitivity of ~ 0.6 eV (90% confidence limit) for Σm_ν .¹⁷ Moreover, the SPTpol constraints compliment other cosmological probes of the neutrino mass (e.g., galaxy clusters), which addresses potential uncertainties due to model dependence, and are comparable to the largest neutrino mass splitting of ~ 0.05 eV.

2.1.3 Simulated Polarization Observations

We estimate SPTpol constraints on the E-mode and B-mode power spectra by performing Monte-Carlo simulations of three years of SPTpol observations that uniformly map a 625 deg^2 field. The simulations assume a realistic scan strategy with 1 deg s^{-1} azimuth scans, and a 50% observing duty cycle during the Austral winter, which is conservative compared to the $\sim 60\%$ achieved by SPT-SZ. The simulations include the effect of foreground contamination, which are mitigated through our choice of a relatively dust-free observing region, see Section 5, and suppressed further by making linear combinations of the two observing bands. The simulations assume an instrumental $1/f$ knee at 0.05 Hz, and further suppression of atmospheric $1/f$ from pair differencing detectors in the same pixel. For this, we assume a frozen-screen atmospheric model and a factor of 200 common-mode rejection from the differencing. In SPT-SZ, we could achieve a factor of 100 reduction in atmospheric signal differencing neighboring detectors without any gain matching, and preliminary analyses of SPTpol indicate better performance.^{18,19} We use a version of the Smith et al. 2006 contaminant-free B-mode estimator²⁰ in these simulations and find the residual E \rightarrow B leakage to be below levels relevant for SPTpol, even when applied to non-cross-linked data that have been filtered to reduce the effects of $1/f$ noise. We fit the B-mode spectra jointly to a gravitational-wave template spectrum and a lensing template spectrum and estimate joint constraints on r and Σm_ν . We note that effects such as polarized sidelobe pickup, beam imperfections, and calibration uncertainties are not included in the simulations, but are expected to be sub-dominant to the statistical errors of this survey due to mitigation techniques of design and strategy outlined in Section 5.

2.2 E-mode and Fine Scale CMB Anisotropies

The primary E-mode polarization signal is generated from the same acoustic oscillations as those that source the temperature fluctuations in the CMB, but is sourced from the velocity field rather than the temperature of the plasma. The E-mode power spectrum (EE) is consequently out of phase with the density fluctuations that power the temperature power spectrum (TT). This means the EE power spectrum carries much the same information as the TT spectrum. Although the E-mode signal is generally at least a factor of 6 weaker than the temperature signal, it has the distinct advantage of suffering from far weaker foreground contaminants (e.g. dusty galaxies, galaxy clusters, galactic sources) that become important at high- ℓ and are difficult to separate in the TT spectrum.²¹ For example, at 150 GHz, the dominant high- ℓ foreground contaminants are dusty galaxies, which are expected to be polarized at only the 1–2% level.²² Therefore, at the limit of systematic uncertainty from foreground contamination dominating statistical uncertainty, the EE spectrum can provide improved constraining power on cosmological parameters sensitive to the primordial high- ℓ spectrum.

The primordial fluctuations at high- ℓ are sensitive to the primordial helium abundance, Y_{He} , through its effect on the epoch of recombination. The density of helium affects the electron density during recombination such that, for a given baryon density, a larger helium abundance will result in a lower electron density during recombination, which results in diffusion damping on larger scales (suppression of high- ℓ acoustic peaks). The effects of helium density on electron density during the epoch of recombination also results in small shifts in the location of the acoustic peaks at high- ℓ .²³ The relatively clean measurement of the primordial high- ℓ spectrum provided by E-modes will allow significant improvements on the measurement of Y_{He} , which leads to an independent measurement of the baryon density ($\Omega_b h^2$) through Big Bang Nucleosynthesis (BBN) and helps break the degeneracy between $\Omega_b h^2$ and n_s , the scalar spectral index.

SPTpol will also expand on much of the fine-scale CMB anisotropy work done by SPT-SZ. SPTpol will probe secondary CMB anisotropy due to the background of lower mass SZ clusters below the detection threshold⁶ and the clustering power of the cosmic infrared background (CIB) in the millimeter.^{7,24} Measurements on these scales will also provide constraints on the evolution of the ionized fraction during the epoch of reionization through the kinetic SZ effect.^{9,25} Comparing the gravitational lensing of the CMB to galaxy surveys or other external mass tracers provides another source of constraint on dark energy and structure formation.^{10,26} Fine scale anisotropies are also sensitive to the expansion rate during recombination and thus to the number of relativistic particle species present at that epoch. This has allowed SPT to place the tightest yet constraints on the number of light particle species beyond the standard three neutrinos⁴ (e.g., sterile neutrinos).

2.3 Clusters of Galaxies

Clusters of galaxies are the largest gravitationally bound objects in the Universe. Their large masses make them a unique cosmological probe sensitive to gravity and the growth of structure on the largest physical scales. As demonstrated by SPT-SZ, a high-resolution SZ cluster survey can uniquely find the most massive clusters in the Universe nearly independently of redshift.⁵ Relative to SPT-SZ, SPTpol will have a factor of ~ 1.3 times lower mass threshold and find a comparable number of clusters, with the expected number of clusters per deg^2 being a factor of ~ 3 times larger. The lower mass threshold will also extend the redshift reach of the cluster survey, and test the effect of dark energy and neutrino mass on the growth of structure at even higher redshifts.

2.4 Dusty Star Forming Galaxies

The SPTpol survey will also discover a significant number of faint extragalactic discrete sources at mm-wavelength. The large areas and high sensitivities of the SPT-SZ and SPTpol surveys are ideal for detecting the brightest mm-wavelength objects in the Universe. These bright mm-wavelength sources tend to be strongly lensed dusty star forming galaxies (DSFGs),¹¹ massive, dusty galaxies that make up a significant fraction of the high-redshift ($z > 1$) component of the cosmic infrared background (CIB), and are crucial to our understanding of galaxy formation. This capability makes the SPTpol survey an important compliment to the high spatial and spectral resolution of ALMA, by identifying the most extreme mm-wavelength bright sources in the Universe. With ALMA, the high-redshift lensed structures will be useful measures of star formation near the end of the epoch of reionization. They will also probe the dark matter distribution in the elliptical galaxy lenses, to characterize mass distributions, substructure abundance, and determine the mass-to-light ratio in elliptical galaxies.

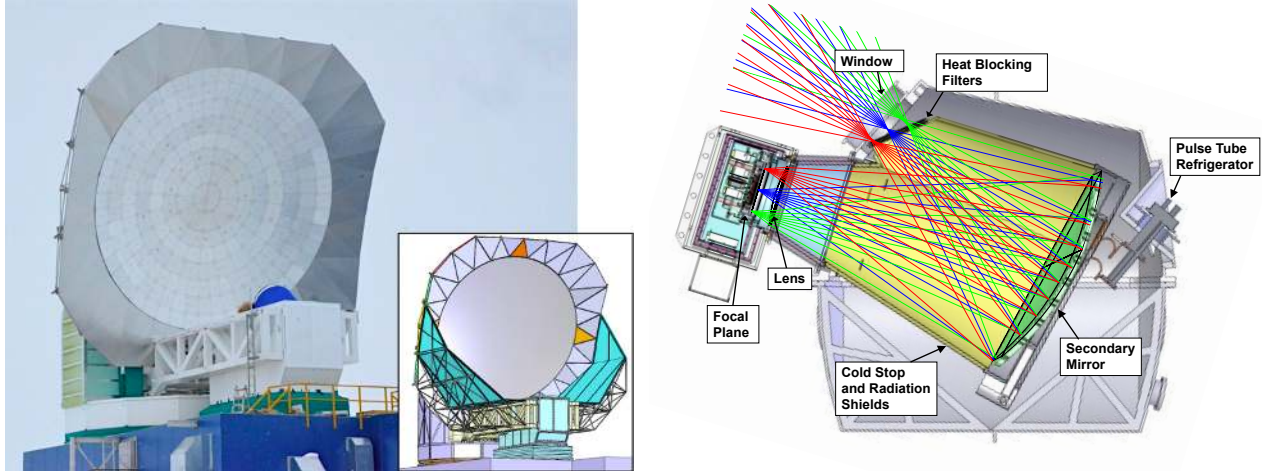


Figure 2. *Left:* Photograph of the SPT with the new RF shield ‘guard ring’ extended past the 10 meter primary, installed during the 2011/12 Austral summer. The inset shows the mechanical design of the second component of new shielding to be installed in November 2012. *Right:* Cross section of the optics and receiver cryostats with representative beams from the center and edges of the focal plane.

3. TELESCOPE AND OPTICS

The SPT is a 10 meter telescope optimized for sensitive, high-angular resolution measurements of the anisotropy of the CMB and mm-wavelength sky.^{27,28} The telescope is located at the NSF Amundsen-Scott South Pole station, the best location on Earth for millimeter-wave observations, with 30 times less atmospheric fluctuation power than found at the ALMA site in the Atacama desert.^{29,30} The telescope is an off-axis, classical Gregorian design that gives a wide diffraction-limited field of view, low scattering and high efficiency with no blockage of the primary aperture. The current telescope optics produce a $1'$ FWHM beamwidth at 150 GHz with a conservative illumination of the inner 8 meters of the telescope, and a $\sim 1 \text{ deg}^2$ diffraction-limited field of view.²⁷ The SPT is designed to modulate the beams on the sky by slewing the entire telescope at up to 4 deg s^{-1} and eliminating the need for a chopping mirror. The telescope operates largely remotely, with a high observing efficiency. In the Austral summer of 2011/2012, the ground shielding of the telescope was improved by adding a $\sim 1.5 \text{ m}$ wide ‘guard-ring’ around the 10 m diameter primary mirror, see Figure 2. In November 2012, additional shielding will be installed along the boom of the telescope for further improvements.

The SPT has a relatively simple optical design with just two mirrors (a primary and secondary) and one lens, with detectors located at the Gregory focus, see Figure 2. From the sky, light enters through a zotefoam vacuum window, which has $>99\%$ transmission averaged across both the 90 and 150 GHz bands. Directly behind the window are a series of infra-red shaders and blockers at ~ 100 and $\sim 10 \text{ K}$, which are tilted to direct reflections away from the focal plane. After the filters, light is reflected off a 1 m secondary mirror, which also effectively acts as an optical stop for the system. For this reason, the secondary is cooled to 10 K with the optical path between the infra-red blocking filters, to the secondary, to the detectors, surrounded by HR-10 microwave absorber also cooled to 10 K. The majority of this optics is housed inside a separate cryostat from the receiver with its own pulse tube refrigerator. Next, another infrared blocking filter sits on top of a lens which re-images the light at the Gregory focus to make a tele-centric focal plane. The lens is made of high density polyethylene (HDPE) with an expanded teflon anti-reflection coating designed to keep reflections at $< 1\%$ averaged across the 90 and 150 GHz bands. The lens is housed in the SPTpol receiver cryostat, and is cooled to $\sim 6 \text{ K}$. Next, a series of low-pass metal-mesh filters are located near the focal plane, to reduce optical loading on the detector cold stage (280 mK), and define the high-end of the detector band-pass.

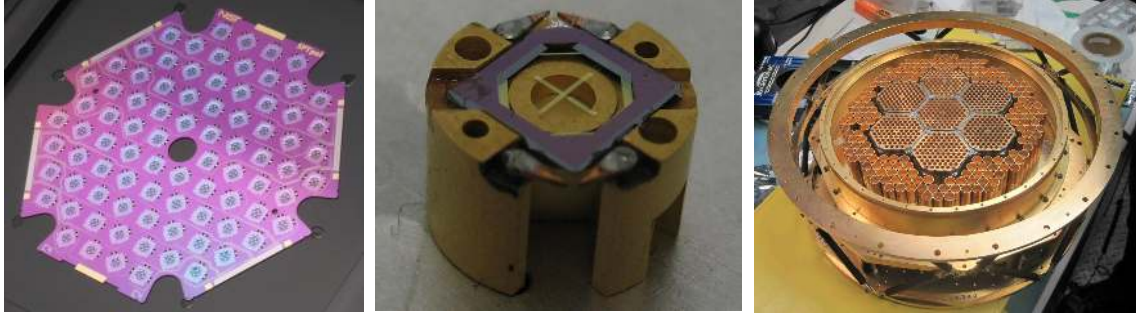


Figure 3. *Left*: Monolithic 150 GHz array, consisting of 84 polarimeters (comprising 168 optical and 5 dark TES bolometers). Pixel orientations vary between 0 and 45° from nominal in order to improve polarization crosslinking when scanning. *Center*: Single 90 GHz pixel in module with feedhorn and waveguide removed. *Right*: Photograph of the SPTpol focal plane populated with 150 GHz array modules (center, with smaller aperture feedhorns) and individual 90 GHz pixel modules (outer rings of feedhorns).

4. SPTPOL RECEIVER

4.1 Focal Plane

The SPTpol focal plane is filled with 768 feedhorn-coupled, dual-polarization pixels, each consisting of two transition-edge sensor (TES) bolometers for simultaneous detection of incident power in each of two orthogonal polarizations (for a total of 1536 optically coupled TES bolometers). The focal plane area is split nearly evenly between 90 and 150 GHz sensitive pixels, with the outer ring of the focal plane consisting of 180 pixels at 90 GHz and the inner diameter consisting of 588 pixels at 150 GHz, see Figure 3. The 90 and 150 GHz pixels were developed independently at Argonne National Laboratory (ANL) and the National Institute of Standards and Technology (NIST) in Boulder, CO, respectively, and are briefly reviewed below with details found elsewhere in these proceedings.^{18,19}

4.1.1 90 GHz Polarimeters

Each 90 GHz pixel is built as a single module that comprises two individual, but identical, single-polarization detectors mounted face-to-face and rotated 90 degrees with respect to each other. A contoured feedhorn with a single-moded circular waveguide couples light to the detectors via a resistive PdAu absorbing bar, which is connected to a SiN thermal mass with a lithographed Mo/Au bilayer TES. In this configuration, the 90 GHz modules have been measured in the lab to have low cross-polarization response ($< 1.6\%$), high optical efficiency for each polarization ($\sim 87\%$), and excellent noise properties consistent with expectations.¹⁸ The 90 GHz modules are built individually and can be installed to the focal plane as independent units, thus allowing replacement of broken pixels or those with poor performance. A 90 GHz pixel with the feedhorn and waveguide removed can be seen in Figure 3.

4.1.2 150 GHz Polarimeters

The SPTpol 150 GHz polarimeter design is based on the generic mm-wave polarimeter development by the TRUCE collaboration³¹⁻³³ that is being used in multiple upcoming and future CMB experiments including ABS³⁴ and ACTpol.³⁵ SPTpol detectors represent a branch of this development where TES designs were modified to match the expected loading and readout requirements of the SPTpol experiment. The layout was expanded into a monolithic array of 84 polarimeters, and on-chip passband filters were removed in favor of low-pass free space filters and a high-pass waveguide in order to improve optical efficiency and block out-of-band leakage.³⁶

Optical power is coupled to the polarimeters through corrugated feedhorns, which help provide low cross-polarization, symmetric beams, high transmission efficiency, and wide bandwidths. Corrugated, gold-plated, silicon platelet feedhorn arrays were developed at NIST³⁷ and made to match the 84-pixel polarimeter arrays. Single-moded waveguide feeds incoming radiation to a planar orthomode transducer (OMT) that separates the radiation into orthogonal polarizations. Each polarization is then coupled through micro-strip to its own TES

Band (GHz)	Number of Pixels (polarimeters)	Resolution (arcmin)	Horn Diam. ($f\lambda$)	Detector & Horn optical eff.	End-to-end optical eff.	Target Survey depth $\mu\text{K-arcmin}$
90	180 (360 TES)	1.6	2.3	~ 0.87	~ 0.46	9
150	588 (1176 TES)	1.0	1.6	~ 0.90	~ 0.45	5

Table 2. Base parameters for the SPTpol polarimeter and projected sensitivity of the full 625 deg^2 SPTpol survey. Each polarimeter comprises two detectors, for a total of 1536 optically coupled detectors in the focal plane. Detector–horn optical efficiency have been confirmed through lab and in situ measurements.^{18,19} System optical efficiency is calculated for a single polarization from a combination of measured and expected efficiencies of all optical components and has been confirmed through observation.³⁸ Depth numbers are projected noise rms (in T; Q and U will be $\sqrt{2}$ higher) in $1'$ pixels for the full 4 years of observation on a 625 deg^2 field and are calculated using achieved mapping speeds and efficiencies of SPT-SZ scaled to the SPTpol sensitivities and yield (Table 3).

bolometer island, where the conductor transitions from superconducting into a lossy gold meander that dissipates the radiation into heat measured by the TES. The 150 GHz passband is defined through a series of low-pass free-space filters located throughout the optical chain and the high-pass cutoff frequency of the waveguide section between the feedhorns and detectors. Lab tests of the feedhorn and detector arrays exhibit excellent cross-polarization (below -25 dB) and optical efficiency ($\sim 90\%$) properties, with measured noise levels consistent with predicted design levels.¹⁹ In addition, the monolithic design results in consistent alignment of pointing and polarization between the two polarization channels of each pixel.

4.2 Readout

All detectors are biased and read out using a second-generation digital frequency-domain multiplexer system with cryogenic SQUID amplifiers based on the initial SPT-SZ readout.^{39,40} This low-noise system introduces a small current-noise term that is sub-dominant to other noise sources under normal operating parameters.³⁸ For SPTpol we multiplex 12 resonant channels on each SQUID,¹⁹ with frequency separation between channels of $\gtrsim 60 \text{ kHz}$ to minimize crosstalk. A total of 144 SQUIDS are used to readout the 1536 optical TES channels in addition to various dark TES bolometers and other calibration channels. Further details on SPTpol readout electronics, control software, data acquisition, data handling and archiving can be found elsewhere in these proceedings.⁴¹

5. OBSERVING STRATEGY

The SPT was designed to conduct large-scale surveys at millimeter and sub-millimeter wavelengths, taking advantage of the exceptionally stable and transparent atmosphere above the South Pole. The unique atmosphere at the South Pole site enables relentless observing in remarkably stable conditions. Furthermore, the geographical location allows the survey observations to be conducted 24 hours-a-day, 7-days-per-week, year-round on a single region of sky through the same airmass: sources do not rise or set. This allows extremely deep, large-area surveys of the southern sky, which includes some of the lowest foreground and most studied CMB fields available (Figure 4), to be completed quickly and with highly uniform sensitivity. Accordingly, all other aspects of the SPTpol project—the telescope, the RF shielding, the receivers, the observing strategy and survey design—have been optimized to minimize systematics and allow ultra-sensitive measurements of the CMB.

5.1 Survey Parameters and Overlap With Other Data Sets

Over a 4 year period that started in February, 2012, SPTpol will map approximately 625 deg^2 to expected depths of ~ 9 and $\sim 5 \mu\text{K-arcmin}$ at 90 and 150 GHz, respectively. The SPTpol survey is located in a region of low galactic foreground and will overlap with the BICEP2 and KECK CMB fields. The exact location and dimensions of the full survey have not been finalized at this time; however, likely locations are depicted in Figure 4.

In advance of the cosmological constraints from the full SPTpol survey, we will leverage previous SPT-SZ data and a wealth of multi-wavelength data to maximize early science from SPTpol in the first year of observation. The entire first year will be spent observing the SPT-SZ 100 deg^2 field centered at right ascension $23^{\text{h}}30^{\text{m}}$, declination -55° , located in one of the most foreground-free regions in the mm-wave sky, see Figure 4. This field is already the deepest large mm-wave field in the sky, having been observed with SPT-SZ to $\sqrt{2}$ below the normal SPT-SZ survey noise levels, and SPTpol will image it to a factor of ~ 3 lower noise in just the first year. The

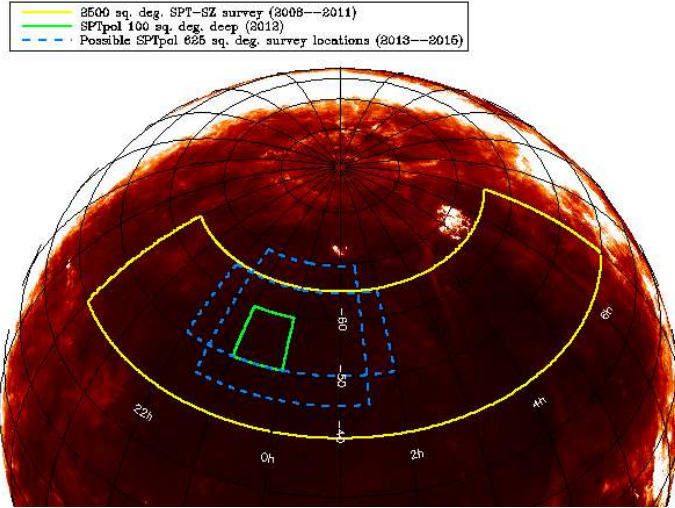


Figure 4. SPTpol survey field locations relative to SPT-SZ plotted over-plotted on an IRAS 100 μm dust map.⁴² The sky is rotated such that the South Celestial Pole is at the top of the globe, and R:A: = 1h faces the viewer

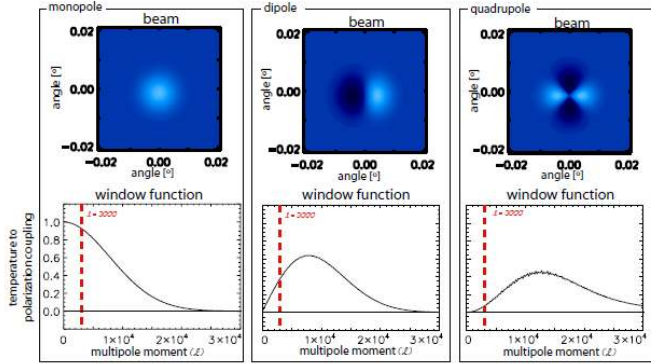
multi-wavelength coverage of this field is unprecedented for a field this size. It has already been covered in the near-infrared and far-infrared/submm with dedicated programs using the *Spitzer*-IRAC and *Herschel*-SPIRE instruments; in the optical, the DES collaboration plans to observe this field to full survey depth in the first months after DECam commissioning; in the X-ray, a large fraction of the field will be covered with a dedicated *XMM-Newton* program; and proposals are being considered to cover the field in several other wavebands.

These multi-wavelength datasets, combined with deep mm-wave temperature and polarization data, will enable several exciting scientific results related to cosmology and the growth of structure in the Universe. The mass limit for SZ cluster detection in this field will be significantly lower than in the SPT-SZ survey, allowing us to extend the SPT cluster catalog to lower mass and higher redshift. This makes the *Spitzer*/IRAC data in this field even more crucial, as measuring redshifts for these clusters will be very difficult in the optical but simple with deep IRAC data.⁴³ The complimentary nature of SPTpol and *Spitzer*/IRAC cluster observations works in the other direction as well: SPTpol will provide mass estimates for *Spitzer*-discovered high-redshift clusters,⁴⁴ as in Brodwin et al. 2012. SPTpol data will similarly provide an initial mass calibration for DES-discovered clusters, and the combination of SPTpol, DES, and *Spitzer*-IRAC data will enable detailed studies of cluster and galaxy formation at high redshift. CMB lensing from the first 100 deg² of SPTpol data will provide high signal-to-noise mass maps of this field, which can be correlated with tracers of large-scale structure from all the multi-wavelength datasets, enabling measurements of galaxy bias at low (DES), intermediate (*Spitzer*), and high (*Herschel*) redshift. In addition, the combination of *Herschel* and SPTpol power spectra will improve constraints on the epoch of reionization.

5.2 Calibration and Mitigation of Polarization Systematics

CMB polarization experiments such as SPTpol build on the deep experience gained from CMB temperature experiments, but the low signal level and the requirement to maintain the fidelity of the polarization pose new challenges. As with most current and planned bolometric CMB polarimeters, SPTpol will extract polarization signals by simultaneous differencing of two detectors sensitive to orthogonal linear polarizations viewing the same point on the sky and portion of the atmosphere. In such a “detector differencing” technique, $T \rightarrow P$ leakage (“instrumental polarization”) can be caused by mismatches in relative gain calibration, beam shape, or bandpass in the two differenced detectors, and by relative pointing errors. Detector polarization angle mis-calibration and pointing reconstruction errors also cause $E \rightarrow B$ leakage (“cross-polarization”). The SPTpol experimental design has numerous features to control polarization systematics, including the large-aperture SPT telescope, a simple, well-shielded optical design, and careful calibration plans.

Temperature to Polarization Leakage Beams and Averaged Window Functions



CMB Power Spectra and T→P Leakage for a 0.017° Beam

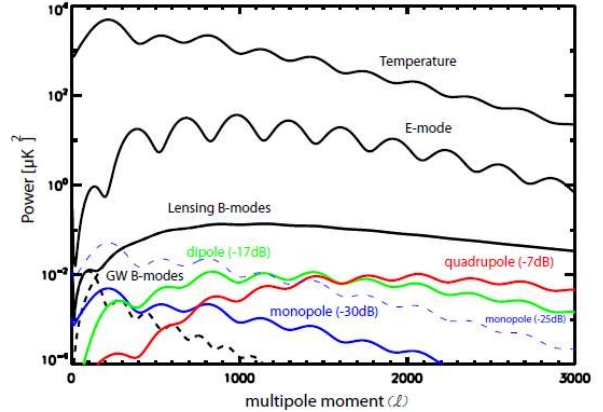


Figure 5. Effects of temperature to polarization ($T \rightarrow P$) leakage for the SPT beam size ($1'$) and the requirements for leakage suppression. *Left*: Beam shapes for monopole, dipole, and quadrupole $T \rightarrow P$ leakage and the corresponding azimuthally averaged window functions, which give the $T \rightarrow P$ leakage as a function of multipole moment ℓ . Note the dipole and quadrupole leakage are highly suppressed by the small SPT beam size where the B-modes peak below $\ell = 3000$. Monopole leakage is accounted for with careful relative calibration, see Section 5.2.3. *Right*: Simulations of the CMB auto-correlation power spectra (Temperature, E-mode, and B-mode) and the leakage power spectra due to monopole (blue), dipole (green), and quadrupole (red) leakage when suppressed by -30 dB, -17 dB, and -7 dB respectively, which keep the $T \rightarrow P$ leakage at least -10 dB below the B-mode signal.

5.2.1 Large-Aperture Telescope

Although the gravitational wave B-mode signal peaks at degree angular scales, the SPT large-aperture telescope, with $1'$ resolution, offers significant advantages in reducing systematic errors generated by beam asymmetries, such as beam shape mismatches in differenced beams, which can cause $T \rightarrow P$ leakage.⁴⁵ As shown in the left panel of Figure 5, the $1'$ SPT beams are sufficiently small that the differenced-beam asymmetries (characterized by beam dipole and quadrupole moments) probe the damping tail of the CMB power spectra where power is low. This attenuation of the $T \rightarrow P$ coupling is a distinct advantage of SPT's small beamsize, relative to other ongoing CMB polarization efforts. Although precision beam characterization of SPTpol awaits the availability of a planet to observe from Pole (e.g. Mars in September 2012), early results and the clean SPTpol corrugated feed design suggest the relatively mild requirements on beam shape mismatch shown in Figure 5 should be easily met.

Another systematic that is mitigated by the high angular resolution of SPT is the requirement on the beam dipole caused by pointing offsets between the two polarization beams from one pixel. Our goal is that the two beams for a pixel be co-located to within about $2''$. This should be achieved given the high precision of the relative pointing of the detectors, and will be verified to a small fraction of the $1'$ beam when planet observations are available.

5.2.2 Simple, Well-shielded Optical Design

The SPT telescope is a very simple design consisting of just two mirrors and one low-power lens. This simple design minimizes spurious polarization generation and distortion. The design obeys the Dragone condition giving zero polarization rotation (crosspol) at the center of the field. The two mirrors and lens give 0.03% $T \rightarrow P$ leakage, which can be accounted for with proper calibration (Section 5.2.3), and optics thermal stability (which is aided enormously by the lack of diurnal variation in the South Pole thermal environment). The telescope is also designed to give high rejection of any emission outside the main beam, achieved with a combination of an off-axis telescope, cold stop at the secondary, and integral co-moving shield.

5.2.3 SPTpol Calibration

Precision measurements of faint polarized sources require special attention to calibration and beam characterization. Absolute gain calibration for SPTpol will be obtained by comparing SPT-SZ CMB temperature anisotropy

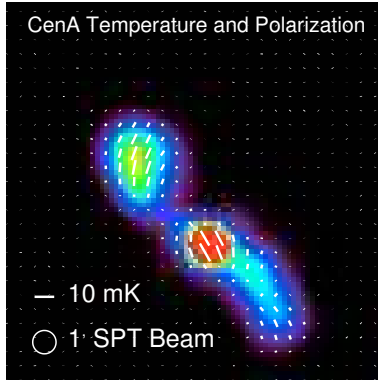


Figure 6. A preliminary image of Centaurus A at 150 GHz showing temperature (false color) and polarization (vectors), obtained during a 40 minute observation of a $\sim 1 \text{ deg}^2$ field during SPTpol commissioning in Feb. 2012.

maps to those produced by *WMAP* and *Planck*. For sky coverage larger than 500 deg^2 , it should be possible to achieve a 1–2% cross-calibration with *WMAP* and likely better with higher-sensitivity and higher-resolution maps from *Planck*. As is being done for SPT-SZ, gain stability will be monitored by periodically activating a chopped IR source viewed through a small hole in the cold secondary, with regularly scheduled elevation “nods” of the telescope that monitor the calibration across the array by observing the zenith-angle dependent atmospheric loading, and with several observations of Galactic sources each day. The response patterns of each detector will be mapped using planets as unpolarized (when unresolved) sources to calibrate the effects of $T \rightarrow P$ leakage in differenced detector pairs.

Absolute orientation angle of the detectors’ polarization response must be calibrated to better than 0.5° to detect an $r = 0.01$ B-mode signal in the presence of the E-modes. The Moon is perhaps the only well-understood, bright, stable and available polarized source in the SPTpol bands with polarization known to this level, but would saturate the detectors to the point where they have little to no response. Although *Planck* and other instruments may continue to improve the catalog of polarization calibration sources*, SPTpol will self-calibrate using a custom-built polarized source located 3 km from the SPT in the far-field of the telescope. The source is a chopped thermal black body polarized by wire grids located in the center of a large reflector that acts to reduce the atmospheric loading on the detectors from the horizon. This system will provide an accurate, high signal-to-noise measurement of the polarization angle of every operational pixel.

6. PROJECTIONS & PERFORMANCE

In Table 3 we give the measured band-pass, and the projected noise and sensitivity of an individual SPTpol TES bolometer during the 2012 observing season (Year 1), predicted using the measured detector parameters, band-passes, and under typical loading conditions. Details of these measurements can be found elsewhere in these proceedings.^{18,19,38} Non-idealities in the Year 1 SPTpol band-passes³⁸ have led to a noise performance that is slightly less than optimal. In Austral summer 2012/2013, the SPTpol filtering will be modified to optimize the high band-edge of both the 90 and 150 GHz bands, and the low band-edge waveguide cutoff of the 150 GHz pixels. These modifications are expected to result in a $\sim 10\%$ improvement in the noise performance, and a $\sim 20\%$ improvement in mapping speed, in each band for observations in 2013 (Year 2) and onward, see Table 3.

The fraction of operational detectors and pixels (yield) during Year 1 was typically $\sim 80\%$. This is expected to be improved in the following years from a combination of replacing cold and warm readout components, and fixing some bad cold wiring. At 80% pixel yield, the Year 2+ combined focal plane NET is projected to be 26 and $15 \mu\text{K} \sqrt{\text{s}}$ for 90 and 150 GHz, respectively. These sensitivities are used for the projected science results from the full 625 deg^2 SPTpol survey, as discussed in Section 2 and summarized in Figure 1 and Table 1. The

*For example, the polarization of Tau A is constrained to better than 1° by combining *WMAP* with other measurements;⁴⁶ however, at a declination of +22, Tau A is not observable from the South Pole

Band	Center Frequency (GHz)	$\Delta\nu$ (GHz)	Optical Power (pW)	NEP _{Total} (aW / $\sqrt{\text{Hz}}$)	NET ($\mu\text{K} \sqrt{\text{s}}$)
YEAR 1					
90	91.2	30.1	10.2	95.5	501
150	146.0	43.4	9.2	76.4	517
YEAR 2+					
90	94.1	37.8	13.9	104.9	442
150	148.5	45.5	8.22	73.1	469

Table 3. Frequency band and noise performance estimates for the SPTpol receiver both before (Year 1) and after (Years 2+) anticipated band-pass upgrades. Year 1 band-pass numbers were measured at the South Pole. Noise estimates are given for an individual TES bolometer, predicted using the measured detector parameters, band-passes, and under typical loading conditions. The predicted noise levels agree to within the 10–20% systematic uncertainties of preliminary on-sky noise measurements.³⁸

projected map depths for SPTpol also agree with a simple scaling from SPT-SZ using its measured focal plane sensitivity and survey map depth, which already include all overheads, data cuts, and observing inefficiencies, and would similarly predict final SPTpol survey depths of ~ 9 and $\sim 5\mu\text{K-arcmin}$ for the 90 and 150 GHz bands, respectively.

ACKNOWLEDGMENTS

Work at the University of Colorado is supported by the NSF through grant AST-0705302. Work at NIST is supported by the NIST Innovations in Measurement Science program. The McGill authors acknowledge funding from the Natural Sciences and Engineering Research Council, Canadian Institute for Advanced Research, and Canada Research Chairs program. MD acknowledges support from an Alfred P. Sloan Research Fellowship. Work at the University of Chicago is supported by grants from the NSF (awards ANT-0638937 and PHY-0114422), the Kavli Foundation, and the Gordon and Betty Moore Foundation. Work at Argonne National Lab is supported by UChicago Argonne, LLC, Operator of Argonne National Laboratory (“Argonne”). Argonne, a U.S. Department of Energy Office of Science Laboratory, is operated under Contract No. DE-AC02-06CH11357. We acknowledge support from the Argonne Center for Nanoscale Materials.

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