

## THE COSMIC BPT DIAGRAM: CONFRONTING THEORY WITH OBSERVATIONS

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### ABSTRACT

We compare a large sample of galaxies between  $0.5 < z < 2.6$  with theoretical predictions for how the optical diagnostic line ratios in galaxy ensembles change as a function of cosmic time. We show that star-forming galaxies at high redshift ( $z > 1.5$ ) are consistent with a model in which the interstellar medium conditions are more extreme at high redshift than seen in the global spectra of local galaxies. We speculate that global spectra of our high-redshift galaxies may be dominated by H II regions similar to the extreme clumpy, dense star-forming complexes in the Antennae and M82. The transition to local-type conditions occurs between  $0.8 < z < 1.5$ . We conclude that classification schemes developed for local samples should not be applied at high redshift ( $z \geq 1.5$ ). We use our theoretical models to derive a new redshift-dependent classification line that utilizes the standard optical diagnostic line ratios  $[\text{O III}]/\text{H}\beta$  and  $[\text{N II}]/\text{H}\alpha$ . Our new line can be used to separate star-forming galaxies from active galactic nuclei (AGN) between  $z = 0$  to  $z \sim 3.5$ . We anticipate that our redshift-dependent optical classification line will be useful for future large surveys with near-infrared multi-object spectrographs. We apply our classification line to a sample of gravitationally lensed galaxies at  $z \sim 2.5$ . Although limited by small numbers, we show that our classification line is consistent with the position of AGN that have been independently confirmed via other methods.

*Key words:* galaxies: abundances – galaxies: active – galaxies: starburst

*Online-only material:* color figures

### 1. INTRODUCTION

Understanding how the fundamental physical properties of ensembles of galaxies change across cosmic time is one of the primary probes of galaxy evolution. Spectral classification is crucial for this work. Classification of galaxies according to their dominant power source is required for measuring the star formation history of galaxies (e.g., Madau et al. 1996; Lilly et al. 1996; Hopkins & Beacom 2006; Sobral et al. 2013, and many others), studying the main sequence for star-forming (SF) galaxies (Noeske et al. 2007), understanding the metallicity history of galaxies (Kobulnicky & Kewley 2004; Zahid et al. 2011; Yuan et al. 2013), and disentangling the cosmic evolution of the starburst versus active galactic nucleus (AGN) fraction (Ramos Almeida et al. 2013).

Baldwin et al. (1981) first proposed the  $[\text{N II}]/\text{H}\alpha$  versus  $[\text{O III}]/\text{H}\beta$  diagnostic diagram (now known as the BPT diagram) to separate normal H II regions, planetary nebulae, and objects photoionized by a harder radiation field. A hard radiation field can be produced by either a power-law continuum from an AGN or from shock excitation. Veilleux & Osterbrock (1987) extended and refined this classification scheme. They used theoretical photoionization models to inform the shape of the classification line between SF and AGN galaxies, and they added two new diagnostic diagrams that exploit the  $[\text{S II}]/\text{H}\alpha$  and  $[\text{O I}]/\text{H}\alpha$  line ratios. Kewley et al. (2001) combined stellar population synthesis and photoionization models to build the first purely theoretical classification scheme for separating pure AGN from galaxies containing star formation.

The Sloan Digital Sky Survey revolutionized optical classification schemes, allowing the shape and position of the local SF abundance sequence and the local starburst-AGN mixing sequence to be cleanly characterized (Kauffmann et al. 2003; Kewley et al. 2006).

Spectral classification beyond  $z > 0.8$  has been difficult in the past. The  $[\text{N II}]/\text{H}\alpha$  ratio is redshifted into the near-infrared (NIR) at  $z \sim 0.5$ , while at redshifts  $z \sim 1.5$ , all of the standard optical diagnostic ratios are in the NIR. Fortunately, NIR multi-object spectrographs on 8–10 m telescopes now allow statistically significant numbers of galaxies to be classified according to their dominant power source.

The  $[\text{N II}]/\text{H}\alpha$  and  $[\text{O III}]/\text{H}\beta$  ratios have now been observed for small numbers of galaxies at high redshift (e.g., Erb et al. 2006; Hainline et al. 2009; Bian et al. 2010; Rigby et al. 2011; Yabe et al. 2012 and references therein). The majority of these galaxies exhibit larger  $[\text{N II}]/\text{H}\alpha$  and  $[\text{O III}]/\text{H}\beta$  ratios than seen in local SF galaxies. Brinchmann et al. (2008) used stellar population synthesis and photoionization models to show that local galaxies with large  $[\text{N II}]/\text{H}\alpha$  and  $[\text{O III}]/\text{H}\beta$  ratios have a larger ionization parameter than the general SF population. They suggest that higher electron densities and a larger escape fraction of H-ionizing photons may be responsible for the larger ionization parameter.

In Kewley et al. (2013), we combined the chemical evolution predictions of cosmological hydrodynamic simulations with stellar evolutionary synthesis and photoionization models to predict how the optical emission-line ratios in SF and AGN populations will change with redshift for four different model

scenarios. We showed that the position of SF galaxies may change significantly as a function of redshift if the interstellar medium (ISM) conditions and/or the stellar ionizing (EUV) radiation field within the SF population changes with cosmic time.

In this Letter, we compare our new theoretical predictions with a statistically significant sample of galaxies between  $0.8 < z < 3$  for which NIR spectra are available. We speculate on the SF conditions in high-redshift galaxies, and we develop a new optical classification scheme that can be applied between  $0 < z < 3.5$ . We adopt the flat  $\Lambda$ -dominated cosmology from the 7 yr *Wilkinson Microwave Anisotropy Probe* experiment ( $h = 0.72$ ,  $\Omega_m = 0.29$ ; Komatsu et al. 2011).

## 2. SAMPLE

Our sample consists of 76 galaxies from magnitude-limited and lensed samples with  $z \gtrsim 0.5$  and measurable [N II], H $\alpha$ , [O III], and H $\beta$  emission lines ( $S/N \geq 3\sigma$ ). These samples contain the most massive ( $M > 10^{9.5} M_\odot$ ) actively SF galaxies at their respective redshifts and may be missing low-mass, low-metallicity galaxies, characterized by low [N II]/H $\alpha$  and high [O III]/H $\beta$  ratios. This selection is sufficient to test the applicability of the local classification scheme at intermediate and high redshift because the most massive, actively SF galaxies lie closest to the starburst–AGN mixing sequence of the BPT diagram (e.g., Yuan et al. 2010). Extending our results to the general population of high-redshift SF galaxies (i.e., fully characterizing the SF abundance sequence with redshift) requires galaxies to be sampled over  $\sim 4$  mag of stellar mass and is not possible with current samples.

At  $z \sim 0.8$ , we use galaxies from the zCOSMOS spectroscopic survey (Lilly et al. 2007, 2009) of the 1.5 deg<sup>2</sup> COSMOS field (Scoville et al. 2007). The zCOSMOS-bright sample contains VIMOS spectra of  $\sim 20,000$  galaxies with  $I_{AB} \leq 22.5$  between  $0 < z < 1.4$ , yielding [O II], H $\beta$ , and [O III] line fluxes for galaxies between  $0.5 < z < 0.92$ . NIR Very Large Telescope ISAAC spectroscopy allowed the additional measurements of H $\alpha$  and [N II] for a sample of 18 galaxies between  $0.5 < z < 0.92$  and two galaxies at  $z \sim 2.5$  (C. Maier et al. 2013, in preparation). For our analysis of the [O III]/H $\beta$  line ratio, we also use the 22 zCOSMOS-bright galaxies with [O III]/H $\beta$  measurements but  $3\sigma$  upper limits on [N II]. This supplementary sample with [N II] upper limits allows us to probe lower metallicity, lower stellar mass ( $10^{9.5} - 10^{10.5} M_\odot$ ) galaxies for our [O III]/H $\beta$  ratio analysis. The zCOSMOS sample excludes AGN identified with X-ray data. We supplement the zCOSMOS data with six galaxies at  $z \sim 0.8$  from the DEEP2 Galaxy Redshift Survey (Davis et al. 2003; Faber et al. 2007) observed by Shapley et al. (2005) and Liu et al. (2008). We also include one gravitationally lensed galaxy observed by Christensen et al. (2012).

At  $z \sim 1.5$ , we use 87 galaxies observed by Yabe et al. (2012, 2013) from the Subaru–XMM Deep Survey (SXDS) (Furusawa et al. 2008) and the UKIDSS Ultra Deep Survey (UDS; Lawrence et al. 2007). The SXDS has limiting magnitudes of  $\sim 27$  in the  $B$ ,  $V$ ,  $R_C$ , and  $i'$  bands and a limiting magnitude of 26 in the  $z'$  band, while the UDS is more shallow (limiting magnitudes of  $\sim 24$ – $25$  in the  $J$ ,  $H$ , and  $K$  bands). Yabe et al. selected galaxies for NIR observations to satisfy:  $K_s < 23.9$  mag, stellar mass  $M_* \geq 10^{9.5} M_\odot$ , and expected H $\alpha$  flux  $F_{H\alpha} > 5.0 \times 10^{-17}$  erg s<sup>-1</sup> cm<sup>-2</sup>. The sample excludes X-ray sources, thus avoiding luminous AGN. Of the 87 Yabe et al. galaxies, 20 objects have  $3\sigma$  detections for all 4 BPT emission lines. The remaining galaxies have upper limits on

either [O III], [N II], or H $\beta$ , which we use for our [O III]/H $\beta$  line analysis. We supplement the Yabe et al. data with six DEEP2 galaxies between  $1.36 \leq z \leq 1.50$  observed by Shapley et al. (2005) and Liu et al. (2008), and six gravitationally lensed galaxies by Rigby et al. (2011), Yuan et al. (2013), and Christensen et al. (2012).

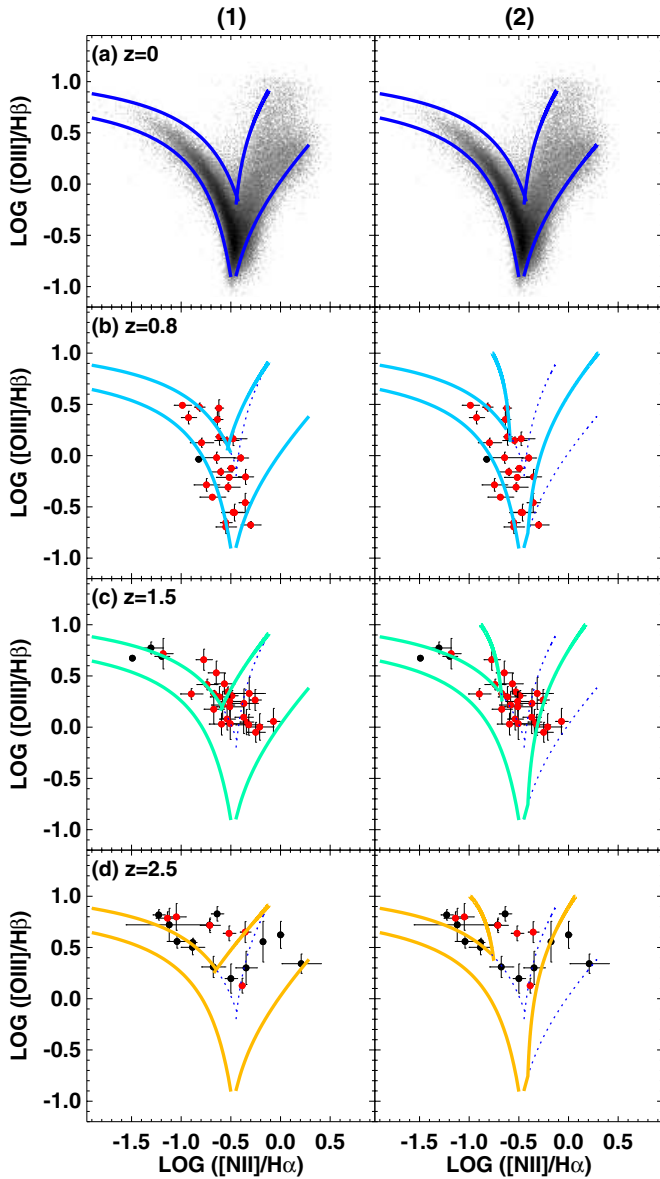
At  $z \sim 2.5$ , we use five galaxies from the SINS sample (Förster Schreiber et al. 2009), two zCOSMOS galaxies (C. Maier et al. 2013, in preparation), the low-metallicity  $L_*$  galaxy (Q2343-BX418) from Erb et al. (2010), and 16 gravitationally lensed galaxies from the compilations of Richard et al. (2011), Yuan et al. (2013), and Jones et al. (2013). Of these, a total of 19 galaxies have  $>3\sigma$  detections of the [N II], H $\alpha$ , [O III], and H $\beta$  emission lines. Gravitational lensing boosts the luminosity of galaxies by 10–30 times, allowing lower luminosity, lower metallicity galaxies to be sampled compared with magnitude-limited surveys. The magnification (and hence the limiting magnitude) differs for each lensing source, depending on the geometry of the lens and the background galaxy. Although the lensing magnification factor does not affect the emission-line ratios, the distortion of the galaxy image causes aperture effects. Hainline et al. (2009) show that this effect is likely to be small (0.05–0.2 dex in the [O III]/H $\beta$  and [N II]/H $\alpha$  line ratios).

The combination of magnitude-limited and lensed galaxies gives a total of 25, 32, and 19 galaxies at  $z = 0.8$ , 1.5, and 2.5 with measured [N II]/H $\alpha$  and [O III]/H $\beta$  line ratios. We also analyze a further 97 galaxies from these samples with upper limits on either [O III], [N II], or H $\beta$ .

## 3. TESTING THE COSMIC BPT DIAGRAM

In Kewley et al. (2013), we combine the chemical evolution predictions from Davé et al. (2011) with the Starburst99 and Pegase2 evolutionary synthesis models (Leitherer et al. 1999; Fioc & Rocca-Volmerange 1999) and our Mappings IV photoionization code (Dopita et al. 2013). Briefly, we use the Dave et al. chemical evolution predictions for the SF gas in  $M_* > 10^9 M_\odot$  galaxies with momentum-conserving winds ( $v_{\text{ZW}}$  in Dave et al.). We use the instantaneous and continuous burst models from Starburst99 and Pegase2 with a Salpeter initial mass function, and the Pauldrach/Hillier stellar atmosphere models. The resulting stellar population spectra are used as the ionizing source in our Mappings IV photoionization code, with a spherical nebular geometry. Mappings IV uses a Kappa temperature distribution which is more suitable for a turbulent ISM than a Stefan–Boltzmann distribution (Nicholls et al. 2012). For AGN, we use the dusty radiation-pressure-dominated models of Groves et al. (2004). Further details of all models used are given in Kewley et al. (2013) and Dopita et al. (2013). Our combination of these models produces theoretical predictions for how the optical emission-line ratios will appear for galaxy samples at different redshifts, based on two limiting assumptions for SF galaxies and AGN.

1. SF galaxies at high redshift ( $z = 3$ ) may have ISM conditions and/or an ionizing radiation field that are either the same as local galaxies (*normal ISM conditions*) or else are more extreme than local galaxies (*extreme ISM conditions*). Extreme conditions in SF galaxies can be produced by a larger ionization parameter (by a factor of 2) and a denser ISM (by a factor of 10), and/or an ionizing radiation field spectral energy distribution (SED) that contains a larger fraction of photons able to ionize O<sup>+</sup>

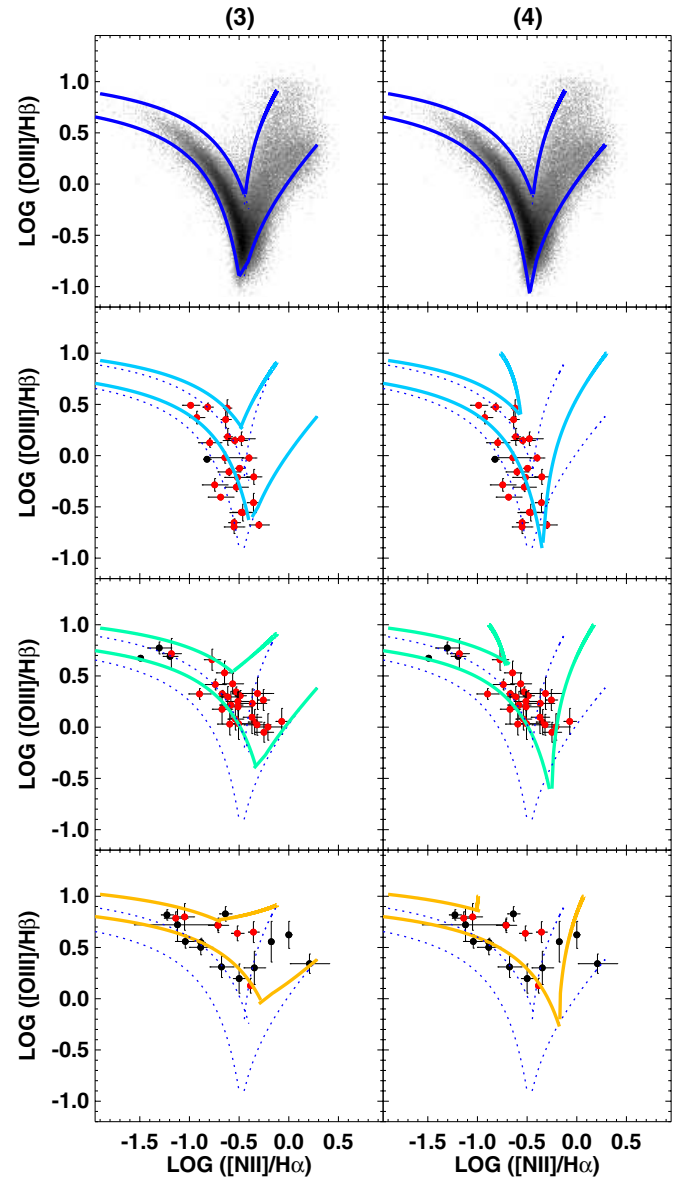


**Figure 1.** The cosmic BPT diagram for four different redshift ranges (rows) and our model scenarios (1) and (2). In each panel, solid lines show our theoretical predictions for the position of the star-forming abundance sequence (left curves) and the AGN mixing sequence (right curves). Black and red circles indicate the positions of gravitationally lensed and magnitude-limited survey galaxies, respectively. Blue dotted lines indicate the boundaries at  $z = 0$ , for comparison. (A color version of this figure is available in the online journal.)

into  $O^{++}$  (i.e., energy  $> 35.12$  eV) relative to the number of H-ionizing photons.

2. The AGN narrow-line region (NLR) at high redshift may either have already reached the level of enrichment observed in local galaxies (*metal-rich*), or else it has the same metallicity as the surrounding SF gas. In the latter case, the AGN NLR at high redshift would be more *metal-poor* than local AGN NLRs. The metal-rich AGN case is equivalent to high- $z$  AGN galaxies containing steeper metallicity gradients than seen in local galaxies, while the metal-poor AGN case is equivalent to high- $z$  AGN galaxies containing flatter metallicity gradients than seen in local galaxies.

These limiting assumptions are described in detail in Kewley et al. (2013). In this work, we focus on the position of the SF



**Figure 2.** As in Figure 1, but for our model scenarios (3) and (4). Data at  $z > 1$  are most consistent with scenarios (3) and (4), in which star-forming galaxies have more extreme ISM conditions at high redshift.

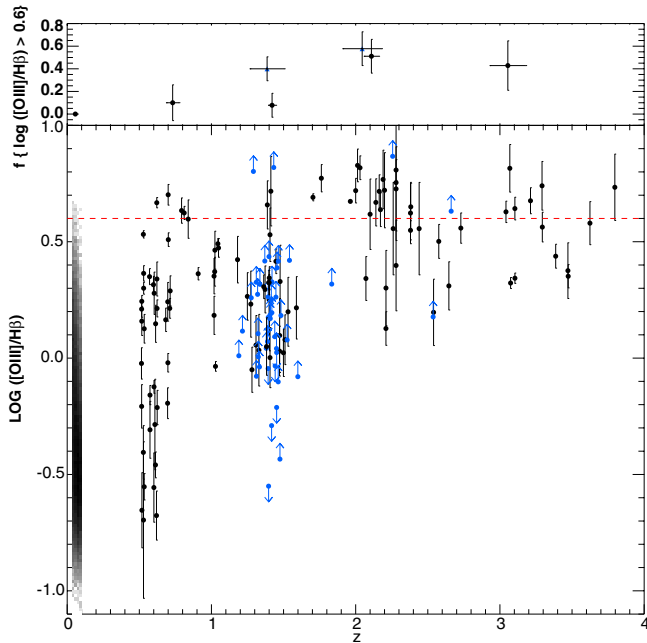
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abundance sequence rather than the position of the AGN mixing sequence.

In Figures 1 and 2, we compare our samples with our model predictions. The columns indicate the four different model scenarios for the SF galaxies and the AGN NLR that are obtained by our two sets of limiting assumptions.

1. Normal SF/ISM conditions; metal-rich AGN NLR at high- $z$  (Figure 1, left panel).
2. Normal SF/ISM conditions; metal-poor AGN NLR at high- $z$  (Figure 1, right panel).
3. Extreme SF/ISM conditions at high- $z$ ; metal-rich AGN NLR at high- $z$  (Figure 2, left panel).
4. Extreme SF/ISM conditions at high- $z$ ; metal-poor AGN NLR at high- $z$  (Figure 2, right panel).

In Figures 1 and 2, galaxies at  $z \sim 0.8$  span a similar range of  $[O\text{III}]/H\beta$  ratios to local SF galaxies. However, at  $z > 1$ , all galaxies in our samples (both lensed and magnitude-limited)



**Figure 3.** Lower panel: the  $[\text{O III}]/\text{H}\beta$  ratio vs. redshift, including  $[\text{O III}]$  and  $\text{H}\beta$  upper limits. Top panel: fraction of galaxies with  $\log([\text{O III}]/\text{H}\beta) > 0.6$  for five redshift intervals. The fraction that would be obtained if all galaxies with  $[\text{O III}]/\text{H}\beta$  lower limits  $\log([\text{O III}]/\text{H}\beta) < 0.6$  truly have  $\log([\text{O III}]/\text{H}\beta) < 0.6$  (black circles) or if they have  $\log([\text{O III}]/\text{H}\beta) > 0.6$  (blue triangles) are shown. The rise in  $\log([\text{O III}]/\text{H}\beta) > 0.6$  with time from  $z \sim 0$  is statistically significant for  $z \gtrsim 2$ .

(A color version of this figure is available in the online journal.)

are offset toward larger  $[\text{O III}]/\text{H}\beta$  ratios than local SF galaxies over the same  $[\text{N II}]/\text{H}\alpha$  range (i.e., the same metallicity range). This result is consistent with previous work on smaller samples (e.g., Shapley et al. 2005; Lehnert et al. 2009; Rigby et al. 2011; Yabe et al. 2012).

We investigate whether this offset is due to observational detection limits in Figure 3 (lower panel), which gives the  $[\text{O III}]/\text{H}\beta$  ratio as a function of redshift. One advantage of this diagram is that we can include galaxies without measurable  $[\text{N II}]$  and  $\text{H}\alpha$ , allowing us to include higher redshift samples. We include the four Lyman break galaxies from Pettini et al. (2001), seven Lyman break galaxies from Maiolino et al. (2008), and three lensed galaxies from Richard et al. (2011) with measured  $[\text{O III}]/\text{H}\beta$  ratios (to the  $3\sigma$  level). Figure 3 also shows galaxies with  $3\sigma$  upper limits on  $[\text{O III}]$  or  $\text{H}\beta$ , where available. The  $[\text{O III}]/\text{H}\beta$  ratios are systematically larger for galaxies at  $z > 1$ , even when upper limits are taken into account. The majority of galaxies (34/47; 87%) with  $[\text{O III}]/\text{H}\beta$  limits have  $[\text{O III}]$  detections but no  $\text{H}\beta$  detections (i.e., *lower limits* on the  $[\text{O III}]/\text{H}\beta$  ratio), strengthening the rise of  $[\text{O III}]/\text{H}\beta$  with redshift.

The  $[\text{O III}]/\text{H}\beta$  upper envelope also rises with redshift such that galaxies have larger  $[\text{O III}]/\text{H}\beta$  ratios on average at  $z > 1$  than galaxies at lower redshifts, over the same  $[\text{N II}]/\text{H}\alpha$  (or metallicity) range. The fraction of galaxies with  $[\text{O III}]/\text{H}\beta > 0.6$  rises with redshift (Figure 3, top panel). This rise is not caused by detection bias because we are able to detect  $[\text{O III}]/\text{H}\beta$  for AGN with  $[\text{O III}]/\text{H}\beta > 0.6$  in the same observational samples at  $z < 1$ .

We propose that the rise in  $[\text{O III}]/\text{H}\beta$  with redshift is caused by either (or a combination of) (1) a larger ionization parameter and a denser ISM at high redshift, and/or an ionizing SED that contains a larger fraction of  $\text{O}^+$ -ionizing to H-ionizing

photons. The larger ionization parameter may be related to the large specific star formation rate observed at high redshift (e.g., Noeske et al. 2007). We note that in the local galaxies M82 and the Antennae, extremely high ionized gas densities and ionization parameters are found in clumpy, dense SF complexes (Smith et al. 2006; Snijders et al. 2007). The densities and ionization parameters measured in these dense complexes are similar to those observed at high redshift (e.g., Rigby et al. 2011 and references therein). Snijders et al. (2007) shows that a geometrical model in which several individual SF clumps are embedded in a giant molecular cloud can reproduce these extreme ISM conditions. We speculate that global spectra of our high-redshift galaxies may be dominated by  $\text{H II}$  regions similar to the clumpy, dense SF regions in the Antennae and M82. It is also possible that the  $\text{H II}$  regions at high redshift are predominantly matter bounded. In this case, the effective ionization parameter from our radiation-bounded models would appear larger (see Kewley et al. 2013, for a discussion). In a follow-up paper (L. J. Kewley et al. in preparation), we investigate these scenarios in detail.

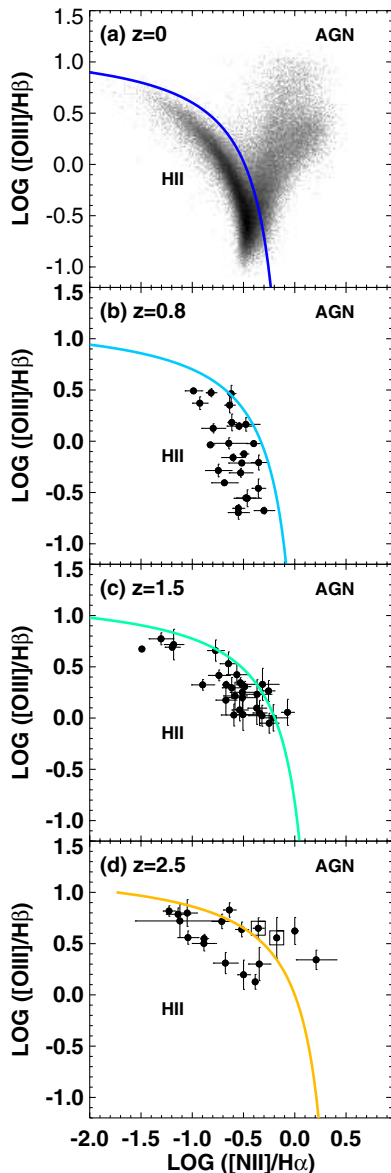
#### 4. NEW REDSHIFT-DEPENDENT CLASSIFICATION SCHEME

Figure 2 indicates that the current optical spectral classification schemes are not suitable for classifying galaxies at  $z > 1$ . SF galaxies are best fit by the theoretical models described in scenarios 3 and 4. In both of these scenarios, the SF abundance sequence can be reproduced by either (1) an ionizing radiation field with a larger fraction of  $\text{O}^+$ -ionizing to H-ionizing photons, and/or (2) a combination of a larger ionization parameter and a high electron density. We use the upper curve of our theoretical abundance sequence to define semi-empirically how the position of the classification line may change with redshift:

$$\log([\text{O III}]/\text{H}\beta) = \frac{0.61}{(\log([\text{N II}]/\text{H}\alpha) - 0.02 - 0.1833 \times z) + 1.2 + 0.03 \times z}, \quad (1)$$

where  $[\text{O III}]/\text{H}\beta$  uses the  $[\text{O III}] \lambda 5007$  line and  $[\text{N II}]/\text{H}\alpha$  uses the  $[\text{N II}] \lambda 6584$  line. Equation (1) can be used to separate purely SF galaxies from galaxies containing AGN, and is based on the following assumptions.

1. The shape of the ionizing radiation field and the ISM conditions in SF galaxies evolves only marginally ( $< 0.1$  dex) with redshift until  $z \sim 1$ , consistent with our findings in Figure 2. This assumption may only hold for galaxies in the mass range of our samples ( $M_* > 10^9 M_\odot$ ). The SF conditions in galaxies with lower stellar masses may still be evolving between  $z = 1$  and  $z = 0$ . If SF conditions are more extreme in low stellar mass galaxies in this redshift range, then these galaxies will lie above and to the right of the line given by Equation (1). In this case, our optical spectral classification line would need to include a stellar mass term.
2. The shape of the ionizing SED and/or the ISM conditions in SF galaxies evolves toward more extreme conditions between  $1 < z < 3$ . Figure 2 suggests that such a change does occur, at least in the most massive galaxies ( $M_* > 10^9 M_\odot$ ). In Figures 1 and 2, we have binned galaxies into two redshift ranges with median redshifts of  $z = 1.5$  and  $z = 2.5$  to provide a sufficient number of galaxies in each redshift interval for statistically significant



**Figure 4.** Our new theoretical redshift-dependent classification scheme for four redshifts. Galaxies at  $z \sim 2.5$  with additional evidence for AGN are shown with a square outline.

(A color version of this figure is available in the online journal.)

conclusions. This small number of redshift intervals allows us to coarsely fit the evolution of the abundance sequence between  $0.8 < z < 2.5$ . A substantially larger sample ( $\gtrsim 200$  galaxies) would allow one to divide the sample into a larger number of small redshift intervals, allowing for a more robust fit in this redshift range.

We note that our chemical evolution assumption affects the position of the intersection between the SF abundance sequence and the AGN mixing sequence, but does not affect the shape or location of the classification line.

In Figure 4 we show the classification line from Equation (1) for our four redshift ranges. The zCOSMOS-bright ( $z \sim 0.8$ ) and Yabe ( $z \sim 1.5$ ) NIR samples excluded bright AGN using X-ray data. The location of the SF galaxies in these samples is consistent with our new classification line, within the errors ( $\pm 0.1$  dex).

Of the five  $z \sim 2.5$  galaxies that we classify optically as AGN, two galaxies (MACS J09012240+1814321 and BzK

15504; square symbols in Figure 4) show additional evidence for an AGN. UV spectroscopy led Diehl et al. (2009) to suggest that MACS J0901 contains an AGN. Although dominant in the UV, the AGN does not make an energetically significant contribution to the global mid-infrared spectrum (Fadely et al. 2010). Lehnert et al. (2009) suggest that BzK15504 contains an AGN based on its strong [O I]/H $\alpha$  emission-line ratio. The remaining three optically classified AGN are A1835 (Richard et al. 2011), MACS J1148 (Jones et al. 2013), and Q2343-BX389 (Lehnert et al. 2009). These three galaxies do not have sufficient ancillary data to confirm or rule out an AGN at other wavelengths. Follow-up of these three galaxies at X-ray, UV, and/or mid-infrared wavelengths will facilitate testing of our new classification scheme.

## 5. CONCLUSIONS

We have presented the first comparison between theoretical predictions of the cosmic evolution of the BPT diagram and NIR spectroscopic observations of active galaxies between  $0.8 < z < 2.5$ . We show that SF galaxies at high redshift are more extreme than a model in which the ISM conditions are more extreme at high redshift. These extreme conditions may manifest in either (or a combination of) a larger ionization parameter, a larger electron density, and/or an ionizing radiation field with a larger fraction of O<sup>+</sup>-ionizing to H-ionizing photons. We speculate that the global spectra of high-redshift galaxies may be dominated by complexes of SF clusters embedded within giant molecular clouds, as seen in the most extreme SF regions in the Antennae and M82.

Due to the change in ISM conditions in SF galaxies with redshift, current optical classification methods based on local samples are not reliable beyond  $z > 1$ . We present a new redshift-dependent optical classification line for the [O III]/H $\beta$  versus [N II]/H $\alpha$  diagnostic diagram that accounts for how the structure of the ISM changes with redshift. We show that the position of independently confirmed AGN is consistent with our classification line at  $z \sim 2.5$ .

In an upcoming paper, we will further investigate the cause of the extreme ISM conditions in high-redshift SF galaxies.

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## REFERENCES

- Baldwin, J. A., Phillips, M. M., & Terlevich, R. 1981, *PASP*, **93**, 5  
 Bian, F., Fan, X., Bechtold, J., et al. 2010, *ApJ*, **725**, 1877  
 Brinchmann, J., Pettini, M., & Charlot, S. 2008, *MNRAS*, **385**, 769  
 Christensen, L., Richard, J., Hjorth, J., et al. 2012, *MNRAS*, **427**, 1953  
 Davé, R., Finlator, K., & Oppenheimer, B. D. 2011, *MNRAS*, **416**, 1354  
 Davis, M., Faber, S. M., Newman, J., et al. 2003, *Proc. SPIE*, **4834**, 161  
 Diehl, H. T., Allam, S. S., Annis, J., et al. 2009, *ApJ*, **707**, 686  
 Dopita, M. A., Sutherland, R. S., Nicholls, D. C., Kewley, L. J., & Vogt, F. P. A. 2013, *ApJ*, submitted  
 Erb, D. K., Pettini, M., Shapley, A. E., et al. 2010, *ApJ*, **719**, 1168  
 Erb, D. K., Shapley, A. E., Pettini, M., et al. 2006, *ApJ*, **644**, 813  
 Faber, S. M., Willmer, C. N. A., Wolf, C., et al. 2007, *ApJ*, **665**, 265  
 Fadely, R., Allam, S. S., Baker, A. J., et al. 2010, *ApJ*, **723**, 729

- Fioc, M., & Rocca-Volmerange, B. 1999, arXiv:astro-ph/9912179
- Förster Schreiber, N. M., Genzel, R., Bouch, N., et al. 2009, *ApJ*, **706**, 1364
- Furusawa, H., Kosugi, G., Akiyama, M., et al. 2008, *ApJS*, **176**, 1
- Groves, B. A., Dopita, M. A., & Sutherland, R. S. 2004, *ApJS*, **153**, 9
- Hainline, K. N., Shapley, A. E., Kornei, K. A., et al. 2009, *ApJ*, **701**, 52
- Hopkins, A. M., & Beacom, J. F. 2006, *ApJ*, **651**, 142
- Jones, T., Ellis, R. S., Richard, J., & Jullo, E. 2013, *ApJ*, **765**, 48
- Kauffmann, G., Heckman, T. M., Tremonti, C., et al. 2003, *MNRAS*, **346**, 1055
- Kewley, L. J., Dopita, M. A., Leitherer, C., et al. 2013, *ApJ*, in press (arXiv:1307.0514)
- Kewley, L. J., Dopita, M. A., Sutherland, R. S., Heisler, C. A., & Trevena, J. 2001, *ApJ*, **556**, 121
- Kewley, L. J., Groves, B., Kauffmann, G., & Heckman, T. 2006, *MNRAS*, **372**, 961
- Kobulnicky, H. A., & Kewley, L. J. 2004, *ApJ*, **617**, 240
- Komatsu, E., Smith, K. M., Dunkley, J., et al. 2011, *ApJS*, **192**, 18
- Lawrence, A., Warren, S. J., Almaini, O., et al. 2007, *MNRAS*, **379**, 1599
- Lehnert, M. D., Nesvadba, N. P. H., Le Tiran, L., et al. 2009, *ApJ*, **699**, 1660
- Leitherer, C., Schaerer, D., Goldader, J. D., et al. 1999, *ApJS*, **123**, 3
- Lilly, S. J., Le Brun, V., Maier, C., et al. 2009, *ApJS*, **184**, 218
- Lilly, S. J., Le Fevre, O., Hammer, F., & Crampton, D. 1996, *ApJL*, **460**, L1
- Lilly, S. J., Le Fèvre, O., Renzini, A., et al. 2007, *ApJS*, **172**, 70
- Liu, X., Shapley, A. E., Coil, A. L., Brinchmann, J., & Ma, C.-P. 2008, *ApJ*, **678**, 758
- Madau, P., Ferguson, H. C., Dickinson, M. E., et al. 1996, *MNRAS*, **283**, 1388
- Maiolino, R., Nagao, T., Grazian, A., et al. 2008, *A&A*, **488**, 463
- Nicholls, D. C., Dopita, M. A., & Sutherland, R. S. 2012, *ApJ*, **752**, 148
- Noeske, K. G., Weiner, B. J., Faber, S. M., et al. 2007, *ApJL*, **660**, L43
- Pettini, M., Shapley, A. E., Steidel, C. C., et al. 2001, *ApJ*, **554**, 981
- Ramos Almeida, C., Rodríguez Espinosa, J. M., Acosta-Pulido, J. A., et al. 2013, *MNRAS*, **429**, 3449
- Richard, J., Jones, T., Ellis, R., et al. 2011, *MNRAS*, **413**, 643
- Rigby, J. R., Wuyts, E., Gladders, M. D., Sharon, K., & Becker, G. D. 2011, *ApJ*, **732**, 59
- Scoville, N., Aussel, H., Brusa, M., et al. 2007, *ApJS*, **172**, 1
- Shapley, A. E., Coil, A. L., Ma, C.-P., & Bundy, K. 2005, *ApJ*, **635**, 1006
- Smith, L. J., Westmoquette, M. S., Gallagher, J. S., et al. 2006, *MNRAS*, **370**, 513
- Snijders, L., Kewley, L. J., & van der Werf, P. P. 2007, *ApJ*, **669**, 269
- Sobral, D., Smail, I., Best, P. N., et al. 2013, *MNRAS*, **428**, 1128
- Veilleux, S., & Osterbrock, D. E. 1987, *ApJS*, **63**, 295
- Yabe, K., Ohta, K., Iwamuro, F., et al. 2012, *PASJ*, **64**, 60
- Yabe, K., Ohta, K., Iwamuro, F., et al. 2013, *MNRAS*, submitted
- Yuan, T.-T., Kewley, L. J., & Richard, J. 2013, *ApJ*, **763**, 9
- Yuan, T.-T., Kewley, L. J., & Sanders, D. B. 2010, *ApJ*, **709**, 884
- Zahid, H. J., Kewley, L. J., & Bresolin, F. 2011, *ApJ*, **730**, 137