

The mass and energy budget of Cassiopeia A

R. Willingale¹, J. A. M. Bleeker², K. J. van der Heyden², and J. S. Kaastra²

¹ Department of Physics and Astronomy, University of Leicester, University Road, Leicester LE1 7RH, UK

² SRON National Institute for Space Research, Sorbonnelaan 2, 3584 CA Utrecht, The Netherlands
e-mail: J.A.M.Bleeker@sron.nl; K.J.van.der.Heyden@sron.nl; j.s.kaastra@sron.nl

Received 11 July 2002 / Accepted 25 October 2002

Abstract. Further analysis of X-ray spectroscopy results (Willingale et al. 2002) recently obtained from the MOS CCD cameras on-board XMM-Newton provides a detailed description of the hot and cool X-ray emitting plasma in Cas A. Measurement of the Doppler broadening of the X-ray emission lines is consistent with the expected ion velocities, $\sim 1500 \text{ km s}^{-1}$ along the line of sight, in the post shock plasma. Assuming a distance of 3.4 kpc, a constant total pressure throughout the remnant and combining the X-ray observations with optical measurements we estimate the total remnant mass as $10 M_{\odot}$ and the total thermal energy as $7 \times 10^{43} \text{ J}$. We derive the differential mass distribution as a function of ionisation age for the hot and cool X-ray emitting components. This distribution is consistent with a hot component dominated by swept up mass heated by the primary shock and a cool component which are ablated clumpy ejecta material which were and are still being heated by interaction with the preheated swept up material. We calculate a balanced mass and energy budget for the supernova explosion giving a grand total of $1.0 \times 10^{44} \text{ J}$ in an ejected mass; approximately $\sim 0.4 M_{\odot}$ of the ejecta were diffuse with an initial rms velocity $\sim 1.5 \times 10^4 \text{ km s}^{-1}$ while the remaining $\sim 1.8 M_{\odot}$ were clumpy with an initial rms velocity of $\sim 2400 \text{ km s}^{-1}$. Using the Doppler velocity measurements of the X-ray spectral lines we can project the mass into spherical coordinates about the remnant. This provides quantitative evidence for mass and energy beaming in the supernova explosion. The mass and energy occupy less than 4.5 sr ($< 40\%$ of the available solid angle) around the remnant and 64% of the mass occurs in two jets within 45 degrees of a jet axis. We calculate a swept up mass of $7.9 M_{\odot}$ in the emitting plasma and estimate that the total mass lost from the progenitor prior to the explosion could be as high as $\sim 20 M_{\odot}$. We suggest that the progenitor was a Wolf-Rayet star that formed a dense nebular shell before the supernova explosion. This shell underwent heating by the primary shock which was energized by the fast diffuse ejecta.

Key words. ISM: supernova remnants – X-rays: ISM

1. Introduction

If we can measure the total mass, the temperature and the bulk velocity of material in a young SNR we can estimate the total energy released by the SN explosion. Coupling this with Doppler measurements we can deproject the mass and energy from the plane of the sky into an angular distribution around the centre of the SN. Here we present further analysis of XMM-Newton data (Willingale et al. 2002) that provides a quantitative assessment of the mass and energy distribution around Cas A. There is a growing body of evidence that the core collapse of massive stars is an asymmetric process. Spectra of supernovae are polarized, neutron stars produced in supernovae have high velocities, mixing of high-Z radioactive material from the core with hydrogen-rich outer layer of ejecta is very rapid, high velocity bullets have been observed in the Vela SNR (Aschenbach et al. 1995) and Cas A itself (Markert et al. 1983; Willingale et al. 2002) is composed of two oppositely directed jets. The analysis presented here confirms the

non-spherical nature of the Cas A SNR and also provides details about the ionization state of the X-ray emitting plasma and the total energy and mass budget of the SN explosion.

2. Composition and dynamics of the plasma

The spectral fit data from Willingale et al. (2002) provide electron temperature kT_e , emission integral $E_I = \int n_e n_H dv$, ionization age $I_t = \int n_e dt$ and elemental abundances for two plasma components (hot and cool) over a grid of 20×20 arcsec pixels covering the face of Cas A. If n_e is the electron density, n_H is the hydrogen density, n_m is the total number density of protons+neutrons (including those bound up in nuclei), n_i is the total number density of ions (nuclei including protons), then, using the elemental abundances and assuming a fully ionised plasma we can calculate the number of electrons per hydrogen atom $R_e = n_e/n_H$, the effective number of protons and neutrons (baryon mass) per hydrogen atom $R_m = n_m/n_H$ and the number of ions per hydrogen atom $R_i = n_i/n_H$. Table 1 summarises these plasma parameters. The mean and rms values

Send offprint requests to: R. Willingale, e-mail: rw@star.le.ac.uk

Table 1. Mean and rms values for those plasma parameters which are not explicitly dependent on the plasma volume, across the face of the remnant. The values are weighted by the shell volume associated with each pixel. R_m/R_i is the mean mass per ion in units of proton mass.

	R_e	R_m	R_i	R_m/R_i
hot	1.23 ± 0.02	1.47 ± 0.05	1.10 ± 0.03	1.33 ± 0.04
cool	289 ± 222	582 ± 449	37 ± 27	15.6 ± 0.7

were calculated by weighting with the shell volume associated with each pixel (see below). The cool plasma component used for the spectral modelling was assumed to be oxygen rich rather than hydrogen rich with all the elemental abundances for the elements heavier than Helium being multiplied by a factor of 10 000. Therefore the electron density and baryon mass per hydrogen atom are high for this component and because of the variations in abundance of the heavy elements there is considerable scatter in R_e , R_m and R_i .

Using the combination of measured Doppler shifts and sky positions for each pixel we were able to estimate the radial distribution of emissivity within the spherical cavity surrounding Cas A (Willingale et al. 2002). The bulk of the emission is confined to a spherical shell of radius 60 to 170 arcsec. Assuming a distance of 3.4 kpc we can calculate the emission volume within the spherical shell associated with each pixel. We know from the high resolution Chandra image of the remnant, Hughes et al. (2000), that the X-ray emission is broken into tight knots and that the plasma doesn't fill the spherical shell. We have therefore assumed filling factors for the components η_{hot} and η_{cool} in each pixel. It is most likely that the emitting material is confined to very irregular, thin, shell-like, filaments which are contained within the larger spherical shell. The filling factors allow for the projection of these structures onto the plane of the sky. The volume of the remnant surrounding these filaments and at smaller radii (<60 arcsec) may contain cold material which has not yet been shock heated. This possibility is discussed further in Sect. 6. Adopting mean values for R_e , R_m and R_i over the plasma volume and as a function of time we can estimate the electron density (n_e), the hydrogen density (n_H), the effective ionisation age (t), the total emitting mass (M) the thermal pressure (P_{th}) and total thermal energy (E_{th}) by using:

$$n_e = \sqrt{E_1 R_e / (V \eta)} \quad (1)$$

$$n_H = \sqrt{E_1 / (V \eta R_e)} \quad (2)$$

$$t = I_t \sqrt{V \eta / (E_1 R_e)} \quad (3)$$

$$M = m_p R_m E_1 t / I_t = m_p R_m \sqrt{E_1 V \eta / R_e} \quad (4)$$

$$P_{\text{th}} = k(T_i R_i + T_e R_e) \sqrt{E_1 / (V \eta R_e)} \quad (5)$$

$$E_{\text{th}} = (3/2)k(T_i R_i + T_e R_e) \sqrt{E_1 V \eta / R_e}. \quad (6)$$

Here m_p is the proton mass, T_i is the ion temperature, V is the total plasma volume and η is a filling factor within that volume.

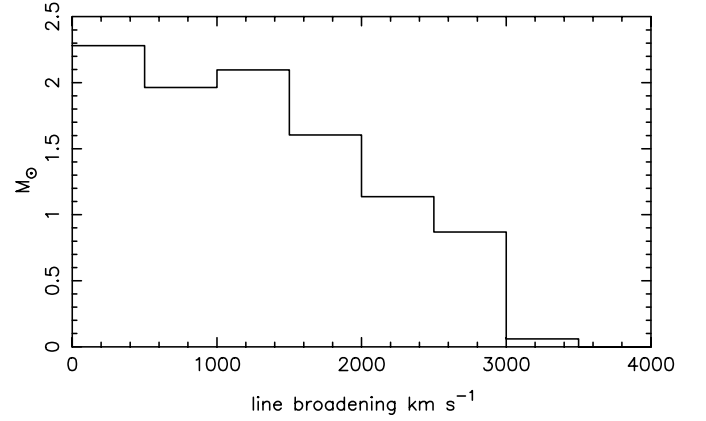


Fig. 1. Mass distribution of line broadening velocity derived from ~ 195 spectral fits across the face of Cas A.

In the spectral fitting we also included a Doppler broadening term to fit the line profiles. Figure 1 shows the mass distribution of the fitted line broadening velocity derived using the mass estimates described below. The rms velocity of the distribution is $1490 \pm 110 \text{ km s}^{-1}$. This represents the rms velocity broadening measured from ~ 195 spectral fits over the face of the remnant. Some of the line broadening may not be due to Doppler but could be introduced by variations in the line blending as a function of temperature which were not accurately modelled using just two temperature components. By looking at the change in line blends over the temperature range of the spectral fits we estimate this introduces a systematic rms error of $\leq 440 \text{ km s}^{-1}$. Estimation of broadening of the line profiles also depends on accurate modelling of the spectral response of the MOS detectors. This is known to an accuracy of a few eV which introduces a possible systematic error of $\pm 500 \text{ km s}^{-1}$. These systematic errors are small compared to statistical error on the individual spectral fits and the distribution of velocities shown in Fig. 1 is dominated by Doppler shift due to the motion of ions in the plasma. The width of the distribution is due to the large spread of ion temperatures (or velocities) within the remnant volume. There are two potential velocity components, thermal motion and micro turbulence. Below we pursue a simple thermal explanation using modelling results from Laming (2001). The spectral fitting gives us a direct measurement of the electron temperature T_e in the plasma but not the ion temperature T_i . Laming (2001) provides predictions of the T_i and T_e for the forward and reverse shocks in Cas A as a function of shock time after the explosion. Using typical ages of the hot and cool components derived below (Table 3) and assuming the hot component is characteristic of the forward shock and the cool component is characteristic of the reverse shock we estimate $(T_i/T_e)_{\text{cool}} = 135$ and $(T_i/T_e)_{\text{hot}} = 10.5$. These ratios are not very sensitive to the ages assumed. The rms thermal velocity along the line of sight is given by $v_{\text{th}} = \sqrt{(kT_i/m_b)}$ where m_b is the mean mass of baryons in the plasma. Therefore using the measured electron temperature, the mean mass per ion listed in Table 1 and T_i/T_e ratios predicted by Laming we can estimate the ion velocity for comparison with the measured Doppler velocities.

Table 2. The measured electron temperature and ion velocity components due to thermal motion (along the line of sight, Doppler broadening) and Doppler shifts of the spectral lines. The velocities are rms mass-weighted values using the mass associated with each pixel.

	T_e keV	v_{th} km s ⁻¹	v_r km s ⁻¹
hot	3.27 ± 0.86	1575	1740
cool	0.45 ± 0.10	608	1780

In Willingale et al. (2002) the Doppler shifts of the prominent emission lines in the X-ray spectrum were used to derive a linear approximation to the radial plasma velocity within the remnant volume,

$$v(r) = \frac{v_p}{(r_s - r_o)}(r - r_o) \quad (7)$$

where the shock radius $r_s = 153''$, the velocity falls to zero at $r_o = 53''$ and the velocity of the plasma just behind the shock is $v_p = 2600 \text{ km s}^{-1}$. We can use this relationship to estimate the rms radial velocity v_r of the plasma components.

Table 2 gives a summary of the ion velocity results. The electron temperatures measured are considerably lower than Laming’s predictions especially for the cool component. This is because the modelling assumes a uniform circumstellar density without clumping. Further modelling by Laming (private communication) confirms that using a $1/r^2$ density dependence in the circumstellar medium reduces the electron temperatures. The high resolution Chandra image (Hughes 2000) indicates a considerable degree of clumping in the plasma which is expected to give rise to higher densities, see Sect. 3. This reduces the equilibration time through Coulomb interactions so the T_i/T_e ratios are probably lower than the values we have adopted. The thermal velocity for the cool plasma is surprisingly low because the mean mass per ion is rather large for this component (very close to pure oxygen as assumed in the Laming 2001 predictions). There is reasonable agreement between the predicted thermal ion velocity and the ion velocity measured from Doppler broadening of the lines indicating that the predicted T_i/T_e ratios are about right. What we actually measure in the spectral fitting is the weighted average of the Doppler broadening from the hot and cool line components combined. This is predicted to be $\sim 1460 \text{ km s}^{-1}$ compared with the measured value of $1490 \pm 110 \text{ km s}^{-1}$. However, some of the Doppler broadening could be due to chaotic motion at large scales rather than microscopic thermal motion of the ions. If this were the case we would require lower T_i values. Chaotic proper motions, over and above linear expansion, have been observed for radio knots, Anderson & Rundick (1995), (see Fig. 6 *ibid*) but they are small compared with the radial component. Fortunately, when calculating the energy associated with these velocities it doesn’t matter whether they are attributable to thermal motion or turbulence and the results in subsequent sections are not dependent on the detailed applicability of the modelling results from Laming (2001).

3. Mass and energy of the plasma

In order to calculate the mass associated with the hot and cool components we must estimate the filling factors. We can do this if we assume that the total pressure is the same in each of the pixels across the face of the remnant and that the cool and hot phases are in pressure equilibrium. The pressure in the plasma has three components, thermal, ram and magnetic and each of these is associated with energy in the remnant. We don’t know the magnetic condition of the hot and cool components but it is reasonable to assume that the magnetic pressure is proportional to the thermal pressure. The magnetic energy is discussed further in Sect. 6. We were unable to detect a large systematic difference between the radial velocities of the two components (see Table 2) and the turbulent velocities are probably small compared with the thermal velocities so the turbulent ram pressure is not important. The ram pressure due to the bulk motion should be comparable to the thermal pressure since $v_r \approx v_{th}$ (see Table 2). We estimate the energy associated with the ram pressure using the bulk expansion velocities. The thermal energy is directly related to the thermal pressure, Eq. (5). We find that a minimum pressure of $7.91 \times 10^{-8} \text{ Pa}$ (N m^{-2}) gives the maximum possible filling factor of $\eta_{hot} + \eta_{cool} \approx 1.0$ peaking near the Western limb of the remnant. The choice of this maximum value of filling factor is a useful reference because it yields the maximum mass and it is consistent with the estimate of the mean filling factor provided by projection of the data in Sect. 4. Using the minimum pressure equilibrium filling factors we have derived values for the electron and baryon densities, ionisation age, total emitting mass and total thermal energy given in Table 3. If we allow the maximum filling factor to drop below 1.0 or we assume a different distance to the remnant the values and ranges in Table 3 will, of course, change.

Table 3. Estimates of the volume filling factors, mean electron and baryon densities, ionisation age, total emitting mass and total thermal energy for the hot and cool components within the spherical shell radius 60–170''. The baryon density n_m is expressed in units of proton mass. The mean and \pm values given are shell volume-weighted. The ionization ages are the median values from the mass ionisation time distribution (see Fig. 3). The ranges quoted for these ages are the 25 and 75 percentiles of the same distribution. The ranges given for the mass and thermal energy are discussed in the text.

	η	n_e cm ⁻³	n_m cm ⁻³	t yrs	M_\odot	E_{th} 10 ⁴³ J
h	0.31	16	19	131	8.31	6.82
	±0.20	±3	±4	101–182	7.42–9.20	6.09–7.55
c	0.009	61	123	80	1.70	0.20
	±0.014	±15	±30	20–273	1.60–1.80	0.19–0.21

We repeated the above analysis using constant values for the filling factors ($\eta_{hot} = 0.31$, $\eta_{cool} = 0.009$) instead of assuming pressure equilibrium. The results were very similar to those in Table 3. The critical factors that effect the results are the the magnitude of the pressure (or filling factor) and the implied large difference between the filling factors for the hot and cool components. Actually the results don’t require strict pressure

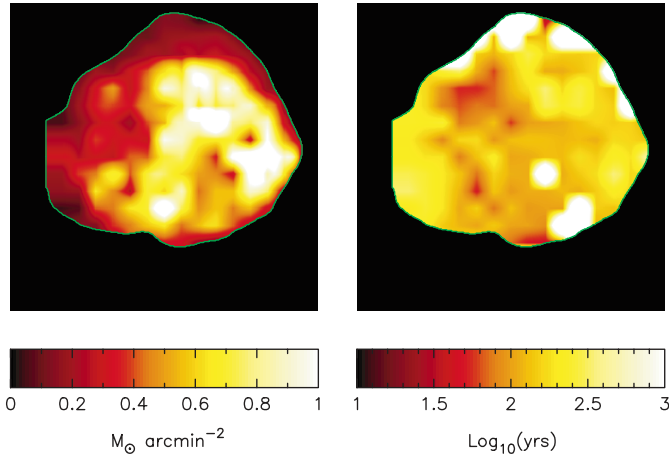


Fig. 2. The maps of mass and ionization age of the hot component. The outer contour line indicates the extent of sky coverage with good statistics for spectral fitting.

equilibrium between the hot and cool components. We simply require the same mean pressure in the two components for a given pixel. The results are also dependent on the volume of the emitting shell. This could be as small as 80–170 arcsec or as large as 50–180 arcsec. If we change the shell parameters the pressure must change in order to give a maximum filling factor of ~ 1.0 . Changing the shell volume within the allowed limits has exactly the same effect as changing the pressure. The corresponding pressure range is $(7.48\text{--}8.34) \times 10^{-8}$ Pa. Using this range we can estimate ranges on the total mass and energy as shown in Table 3.

The maps of the mass distribution and ionization age of the hot component are shown in Fig. 2. The larger ages tend to lie around the perimeter while the ages across the central region are relatively constant with a minimum of ~ 45 years. Figure 3 shows the distribution of the shell volume in the temperature ($T_{\text{hot}}:T_{\text{cool}}$), electron density ($n_{\text{e hot}}:n_{\text{e cool}}$) and ionisation age ($t_{\text{hot}}:t_{\text{cool}}$) planes. The electron density and to some extent the temperatures are correlated over the shell volume. However the ionisation ages are not. Note that the ionisation ages are plotted on logarithmic axes. The spread of age is much larger for the cool component than the hot. The differential mass distribution as a function of ionization age for the two components is also shown in Fig. 3. In this distribution we do see a marked difference between the two plasma components. The hot component shows a relatively sharp peak at an ionization age of ~ 100 years, whereas the cool component has a broader distribution with a median age of ~ 80 years. These profiles indicate that the hot plasma was shock heated over a century ago and the heating process is already complete. The cool plasma, however, has been shocked more recently and the heating process is still ongoing. A small fraction of the mass in both components has an ionization age greater than the true age of the remnant, 320 years, because we have assumed that the densities are constant with time. This is clearly not the case especially when we extrapolate back to the early stages of the remnant expansion.

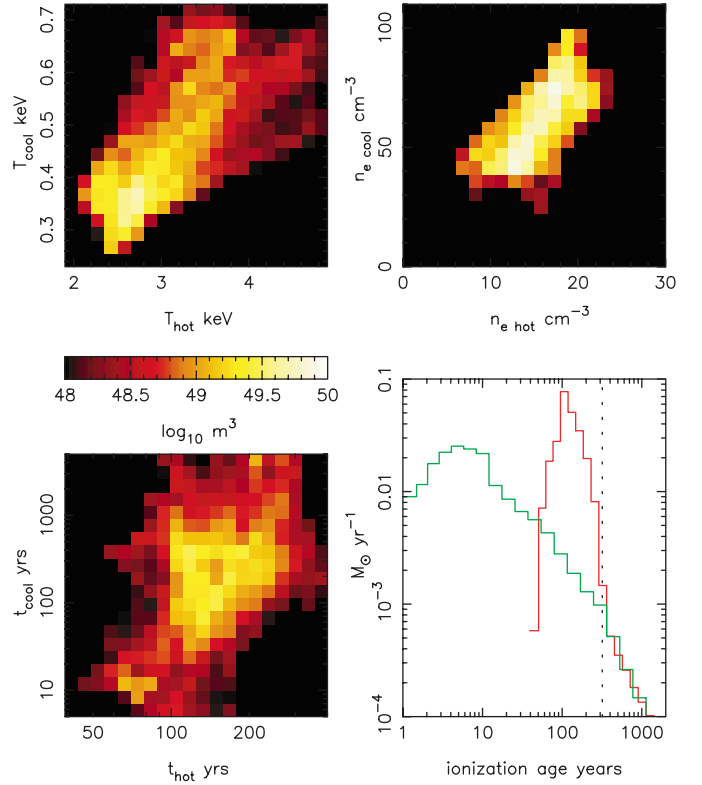


Fig. 3. The distribution of shell volume in the temperature, electron density and ionisation age planes. Note that the volume is plotted on a logarithmic scale to reveal the regions with very low volume. The ionisation age scales are logarithmic showing the spread of the volume over several orders of magnitude of the ionisation time. The bottom right panel shows the differential mass distribution of ionization ages. The hot component is shown in red and the cool component is shown in green. The vertical dotted line indicates the age of the remnant (320 years).

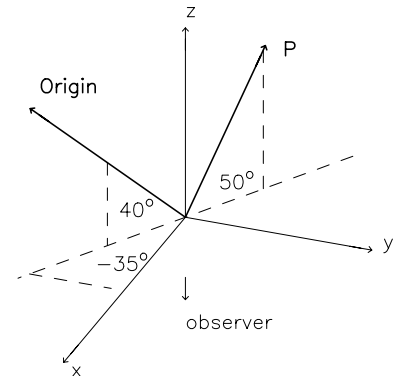


Fig. 4. The *axial* spherical coordinate system with respect to the plane of the sky, x East and y North, and the line of sight z .

4. Projection about the SNR centre

Using the pixel positions with respect to the centre of the remnant on the plane of the sky and the line-of-sight positions provided by the Doppler measurements we can determine the angular distribution of the mass about the centre of the remnant. With x East, y North and z pointing away from the observer we define an *axial* spherical coordinate system as shown in Fig. 4.

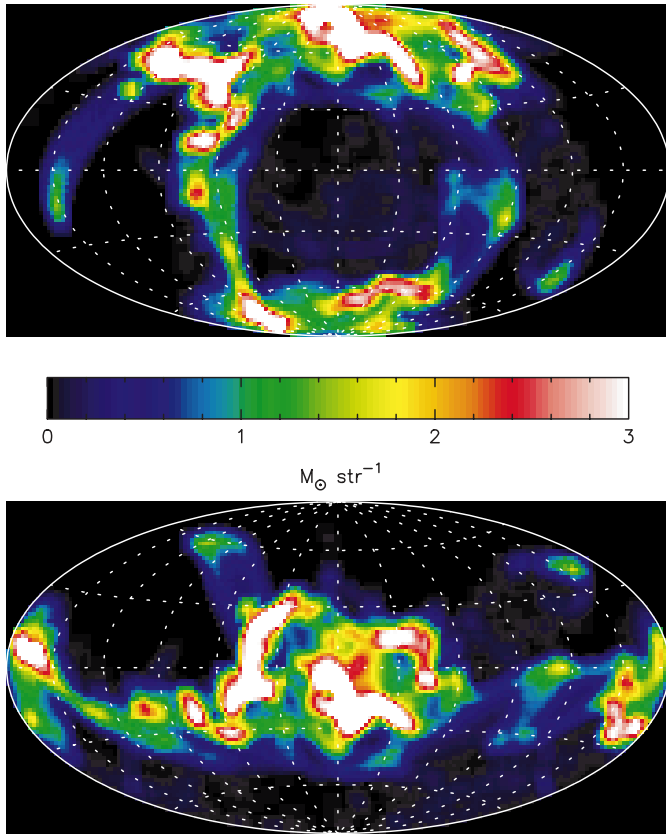


Fig. 5. The angular distribution of mass surrounding the centre of Cas A – Aitoff equal area projection. Upper panel in the *axial* projection and lower panel in the *equatorial* projection.

The polar axis is labelled P and points towards the receding mass in the North of the remnant. The origin on the equator is in the South East quadrant away from the observer. The upper panel of Fig. 5 shows the mass distribution projected in this *axial* coordinate system in Aitoff projection. It is clear that the entire angular distribution of the emitting mass lies in a band around the remnant with enhancements at the poles in the North and South. (As has been suggested by many observations in the past, see for example Markert et al. 1983.) The band of mass is relatively weak when it crosses the equator and there are large solid angle areas around the origin and the anti-centre on the equator where there is very little mass.

The total mass in the Southern hemisphere is $3.3 M_{\odot}$ and in the Northern hemisphere is $6.7 M_{\odot}$ so the split between the two hemispheres is not equal. This simply comes about because there is more emission in the North compared with the South. The asymmetry is significant under the assumptions we have adopted but we have no deeper explanation for it.

Half the mass in the South is contained in 0.63 steradians, a sky fraction of 0.10 and 90% is contained in 2.0 steradians (fraction 0.32). Half the mass in the North is contained in 0.75 steradians, a fraction of 0.12 and 90% is contained in 2.5 steradians (0.39). The left-hand panel of Fig. 6 shows the mass per steradian as a function of elevation angle in the *axial* coordinate system. 64% of the mass is contained within a double cone of half angle 45° and the mass density peaks at

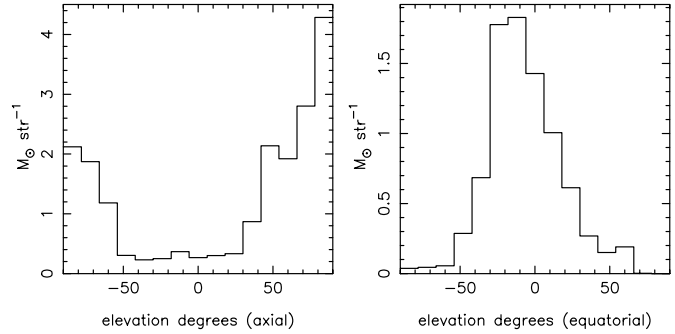


Fig. 6. The mass per steradian as a function of elevation angle. Left-hand panel in the *axial* projection, South-ve elevation, North +ve elevation. Right-hand panel in the *equatorial* projection.

the poles. The jet of high velocity optical material identified by Fesen et al. (1988) and radio knots 89, 90, 92 and 93 listed by Anderson & Rudnick (1995) may be associated with the X-ray axis identified here. This is discussed further in Willingale et al. (2002).

The mass distribution was reprojected into an *equatorial* spherical coordinate system in which the equator lies around the band of mass seen in the top panel of Fig. 5. The North pole was shifted to lie at the origin on the equator. The lower panel in Fig. 5 shows this new projection. The right-hand panel of Fig. 6 shows the distribution about the equatorial plane. In this reprojection a fraction of 0.85 of the mass is confined to within $\pm 30^{\circ}$ of the equatorial plane.

The concentration of mass into a small fraction of the available solid angle as shown in Figs. 5 and 6 is consistent with the mean filling factor of $\eta_{\text{hot}} = 0.31$ (see Table 3) derived from the minimum equilibrium pressure corresponding to a maximum filling factor of ≈ 1.0 . Pressure equilibrium dictates that the filling factor for the cool component is very low, $\eta_{\text{cool}} = 0.009$. This is consistent with the original rationale for the spectral modelling in which the hot component is assumed to be dominated by surrounding medium heated by the primary shock and potentially enriched by diffuse ejecta, whereas the cool component is assumed to be clumpy ejecta heated by the reverse shock. The differential mass distributions with ionisation age are also consistent with this picture. The bulk of the hot component forms a peak with an ionisation age in the range 100–180 years while the cool component has a much broader distribution stretching back to < 10 years, probably indicative of a heating process which is still in progress.

5. The mass and energy budget

The maximum total X-ray emitting mass consistent with the data assuming pressure equilibrium within the remnant volume is $10.0 \pm 0.7 M_{\odot}$. The maximum total thermal energy visible is $7.0 \pm 0.5 \times 10^{43}$ J which is a sizeable fraction of the typical total energy released from a SN explosion. The remainder of the energy in the remnant is kinetic. We can estimate this kinetic energy using the expansion velocities discussed above. To estimate the total mass and kinetic energy in the optical FMKs we have assumed a generous hydrogen density of 10^3 cm^{-3} , (probably reasonably consistent with pressure

Table 4. The mass and energy budget for the remnant. E_{th} is the thermal energy and E_r is the radial expansion kinetic energy. The hot Ejecta entry (2) are the heavy elements in the hot component, while the hot CSM entry (1) is the fraction of the hot component which is not diffuse ejecta and is presumed to have been circumstellar material at the time of the explosion. The cool ejecta (3) are from the cool X-ray component and the cold ejecta (4) are the optical component. The hot diffuse entry refers to the hot component, which consists of diffuse ejecta and swept-up circumstellar matter. The cool clumpy is the sum of the cool X-ray and optical components also assumed to be ejecta. The total is the sum of all the components.

	v_r km s ⁻¹	M_\odot	E_{th} 10 ⁴³ J	E_r 10 ⁴³ J	E_{tot} 10 ⁴³ J
1 hot CSM:	1740	7.9	6.5	2.4	8.9
2 hot ejecta	1740	0.4	0.3	0.1	0.4
3 cool ejecta	1780	1.7	0.2	0.5	0.7
4 cold ejecta	5290	0.1	0.0	0.3	0.3
Σ_1^2 hot diffuse	–	8.3	6.8	2.5	9.3
Σ_3^4 cool clumpy	–	1.8	0.2	0.8	1.0
Σ_2^4 ejecta	–	2.2	0.5	0.9	1.5
Σ_1^4 total	–	10.1	7.0	3.3	10.3

equilibrium) a total of 120 knots, a knot size of 2'' and a velocity of 5290 km s⁻¹ (see Reed et al. 1995; Anderson & Rudnick 1995). Since we have no information about chaotic velocities of optical knots we have assumed these to be zero. Table 4 summarises the resulting mass and energy budget for the remnant. The top half of table is the remnant as we see it now and summing up all the energy components gives an estimate of the total energy released by the explosion, 10.3×10^{43} J. The cool X-ray component, the optical knots and the heavy elements in the hot X-ray component are almost certainly all remains of the ejecta. The evidence for the X-ray components being emission from ejecta material was put forward by Willingale et al. (2002). This evidence includes: i) the elemental abundances of Si, S, Ar and Ca are strongly correlated, ii) the elemental abundance values are consistent with enrichment from ejecta material iii) the emitting material is non-uniformly distributed across the remnant. In particular the emission of Fe-K relative to the lighter elements indicate a large degree of non-uniform mixing. The mass labelled hot ejecta in Table 4 are the heavy elements (not H+He) of the hot component. The presence of these elements in the hot component indicates considerable mixing between the swept up interstellar material and the diffuse ejecta. Previous authors, for example Anderson & Rudnick (1995), have identified this component as diffuse ejecta and estimated the mass as $\sim 0.3 M_\odot$ (Braun 1987), very similar to our estimate of $0.4 M_\odot$. The remaining mass of $7.9 M_\odot$ in the hot component is swept up material labelled hot CSM in Table 4.

Summing up the mass from all the ejecta components indicated in Table 4 and assuming all the energy was kinetic we get the predicted rms ejecta velocity of 6850 km s⁻¹. The original kinetic energy from the diffuse (now hot) ejecta constitutes the driving mechanism of the primary blast wave. Assuming that all the energy in the hot diffuse medium (i.e. 9.3×10^{43} J) was kinetic would require the hot ejecta ($0.4 M_\odot$) to have an initial velocity of $\sim 1.5 \times 10^4$ km s⁻¹. We discuss this (quite

feasible) velocity further in Sect. 6. The original rms velocity of the clumpy (now cool/cold) ejecta must have been much lower, ~ 2400 km s⁻¹. The velocity of the optical FMFs is 8820 km s⁻¹ with a deceleration parameter of 0.99 (the ratio of the true age over the expansion age, $t_{\text{age}}V/R$, where V is the radial proper motion and R is the radius) and the velocity of the optical FMFs is 5290 with a deceleration parameter of 0.98 (see tabulation in Willingale et al. 2002 and references ibid) so these ejecta velocity estimate are entirely reasonable. Estimation of the mass and energy in two identifiable parts of the ejecta gives us the first observational glimpse at the ejecta structure function which plays an important role in analytical and numerical modelling of the early stages of the evolution of SNR, see Truelove & McKee (1999). The values in Table 4 are subject to uncertainties which will only be resolved by observations which much higher spectral resolution but overall the mass and energy budget balances reasonably well.

The hot and cool components also contain 0.012 and $0.046 M_\odot$ of iron respectively. It is reasonable to assume that almost all of this iron originated in the ejecta rather than from swept up interstellar medium since most of the material surrounding the star prior to the explosion probably came from the outer hydrogen rich layers of the progenitor (see discussion below). If this is the case $\sim 2.7\%$ of the diffuse and clumpy ejecta mass was iron, now seen as Fe K emission from the hot component and as Fe L emission from the cool component. The hot iron has a significantly larger radial velocity, 2000 km s⁻¹, than the cool iron, 1580 km s⁻¹, and is seen at larger radii. It is surprising that iron is seen in the diffuse ejecta especially at large radii in the remnant ahead of the lighter elements since it presumably originated from the core of the progenitor not the outer layers of the star. A great deal of mixing of the layered structure of the progenitor must have occurred. This may be because the inner layers were ejected at higher initial velocity than the outer layers and this, in turn, resulted in significant turbulence.

6. Discussion

Just how robust are the values in Table 4? Greatest uncertainty lies in the measurement of the ion temperature and estimation of the volume filling factors. We have set the maximum $\eta_{\text{hot}} + \eta_{\text{cool}} = 1$ while the mean value is ~ 0.3 which is entirely consistent with the observed angular coverage shown in Fig. 5. If we abandon pressure equilibrium the η_{cool} could increase but the η_{hot} would have to decrease and/or the overall filling factor would have to fall. The mean ion temperature is constrained by the measured Doppler broadening of the emission lines. The cool ions could be hotter raising the cool pressure and introducing a pro rata increase in η_{cool} . This would increase the ejecta mass and decrease the swept up mass but the total mass and energy would remain approximately the same. This, in turn, would decrease the rms velocity of the ejecta which at present is consistent with the measured expansion velocities of the optical knots.

We have not included the magnetic pressure (or energy) in the calculations. The electron pressure in the hot component is 4.3×10^{-9} Pa, only $\sim 5\%$ of the total pressure and 7.6×10^{-9} Pa,

$\sim 10\%$ of the total pressure in the cool component. A magnetic field of ~ 2.9 mG will give the same energy density as the electrons in the hot component but such a large equipartition field is unlikely since the field is being amplified by turbulence and magnetic coupling in the post shock plasma. The mean magnetic field required assuming equipartition with the high energy electrons responsible for the radio synchrotron emission is ~ 0.5 mG (Rosenberg 1970; Longair 1994). We conclude that the magnetic energy and relativistic electron energy are only a minor perturbation on the overall energy budget.

Very little of the mass and energy in Table 4 is associated with the faint primary shock which is visible in the Chandra X-ray image (Hughes 2000) and radio images (Anderson & Rudnick 1995). Analysing the Chandra image we find only $\sim 12\%$ of the X-ray flux lies outside the main ring of emission in a region that could be directly associated with the primary shock. We see no distinct spectroscopic signature associated with the primary shock. For example the hard tail of the spectrum is distributed throughout the volume (see Bleeker et al. 2001). If we confine our analysis to the outer filaments in the North and the South, which are allegedly part of the primary shock, we get the same temperatures and densities as given in Tables 1–3, within the quoted errors. It may be that the mass and energy of the primary shock are invisible because the electron temperature has not yet reached the threshold required for X-ray emission. However, we think this is unlikely since the modelling of Laming (2001) indicates that the electrons should reach a temperature of $\sim 3 \times 10^7$ K only 10 years after being shocked and this translates to an angular shift of $2.5''$ on the sky for a shock moving at 4000 km s $^{-1}$. It is possible that slow ejecta lie inside the X-ray emitting shell and this material will remain invisible until it is enveloped by the reverse shock. It is also possible that dense ejecta, which has not been shock heated to a sufficiently high temperature, is hidden in the apparent voids surrounding the optical and X-ray filaments. The existence of such hidden mass could only be inferred from detailed modelling which is beyond the scope of this paper. Any hidden or invisible material would increase the ejecta mass estimate but would have only a minor effect on the total energy. We conclude that hidden mass or energy are unlikely to increase the budget in Table 4 by more than a few percent.

The swept up mass is only seen over about 40% of the total volume of the remnant, Fig. 5. Using an outer radius of $160''$ the implied density of the ambient medium within this volume before the explosion was ~ 13 cm $^{-3}$ which is higher than previous estimates inferred from H II emission (8 cm $^{-3}$, Peimbert 1971) deceleration of the radio-emitting material (2 cm $^{-2}$, Braun 1987) or the low temperature of the X-ray emitting ring (McKee 1974). This high density CSM provides a link between Cas A and a Wolf-Rayet progenitor that suffered mass-loss forming a nebular shell prior to the supernova explosion. Nebular shells around Wolf-Rayet stars are well observed phenomena. They typically have radii of the order $R \sim 1\text{--}4$ pc with shell thickness $\Delta R \sim 0.01\text{--}0.2$ pc and electron densities n_e of a few tens to a few hundreds cm $^{-3}$ (Esteban et al. 1993). We suggest some fraction of a nebula shell was initially heated by the primary shock to form the hot CSM entry in Table 4. The hot ejecta have been mixed with this material by turbulence.

From the difference between the assumed age of the remnant and the time since the the gas was shocked (see the ionisation age in Table 3), we estimate that the system was in free expansion for the first ~ 100 years. This time, coupled with the initial velocity (1.5×10^4 km s $^{-1}$) derived in Sect. 5, implies that the material that we now see as hot ejecta had travelled out to a distance of ~ 1.5 pc ($90''$) before it hit, and heated, the putative dense nebular shell shed by the progenitor WR star. It then swept up about $8 M_{\odot}$ and is now entering the Sedov phase. These dimensions are consistent with both the observed size of the emitting shell and the typical size of nebular shells observed around WR stars. The cooler ejecta, which have a much lower initial velocity, impacted later and show a broader distribution of ionisation ages centered around a lower average absolute value (see Fig. 3). The emission we see from this component is ablated material formed by “reverse” shock-heating of cool, clumpy ejecta.

The total *visible* emitting mass calculated above is lower than previous estimates, Fabian et al. (1980) $15 M_{\odot}$, Vink et al. (1996) $14 M_{\odot}$. The difference is largely attributable to the lower volume estimates. With better spatial resolution and the benefit of Doppler measurements the estimate of the total fraction of the shell volume which is occupied by the emitting plasma is considerably reduced. The spatially resolved high resolution X-ray spectra provided by XMM-Newton also give us a detailed inventory of the state and composition of the plasma which also reduces the uncertainty in estimating the masses involved. However, the mass loss from the progenitor could still be $\sim 20 M_{\odot}$ indicative of a very high loss rate prior to the explosion although only 40% of this material is actually visible.

7. Concluding remarks

We have assumed pressure equilibrium between the hot and cool plasma components to give an estimate of the filling factors within the shell volume. X-ray spectroscopy at higher spectral and spatial resolution could be used to test this assumption. Well resolved emission lines from individual knots would be associated with either the hot or cool component and observation of Doppler broadening of such lines would give us a direct measurement of the ion velocities for individual ion species in the ejecta and swept up material.

Possible supernova core collapse geometries are shown in Fig. 7. The distribution of X-ray emitting mass around Cas A indicates that the original explosion was not symmetric but somewhere between an axial jet and equatorial plane geometry. The confinement to within $\pm 30^{\circ}$ of the equatorial plane as shown in Fig. 6 is rather striking and the other panel in Fig. 6 clearly demonstrates the enhancement of the emission around

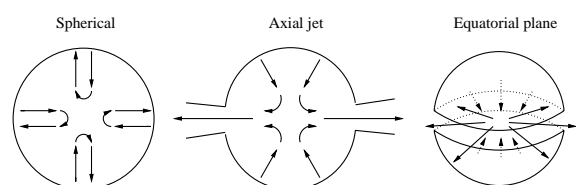


Fig. 7. Supernova collapse geometries.

the poles in the axial coordinate system. It is noteworthy that spherical collapse can be modelled in one dimension, and the axial or equatorial symmetry can be modelled using just two dimensions but the combination of axial and equatorial would require a full three dimensional treatment. It may be that the processes responsible for what we observe will only be revealed by such three dimensional modelling. The apparent asymmetry of the explosion geometry introduces the possibility of significant shear within the expanding material during or just after the explosion. This may be the root cause of the turbulence and the clumpiness of the mass distribution in the remnant rather than hydrodynamic instabilities in the dense shell formed much later after a significant mass of surrounding material has been swept up.

The total kinetic energy derived for the ejecta is consistent with the canonical value of 10^{44} J. The present analysis indicates that the ejected mass was rather large and the rms ejection velocity was correspondingly modest. This is in stark contrast with, for example, the Crab Nebula in which no significant ejected mass or energy from the original explosion has been measured. Hester et al. (1995) claim that the outer rim of optical line emission in the Crab corresponds to the shock advancing through ejecta but they provide no estimate of the ejecta mass or energy involved. Collimated or jet-induced hypernovae have been suggested as a possible solution to the energy budget problem posed by gamma ray bursts seen from cosmological distances, Wang & Wheeler (1998). However in these cases we are looking at collimation in radiation and there is no reason to suppose that the degree of mass collimation seen in Cas A is connected with the radiation collimation inferred in gamma ray burst events although some models for GRBs do involve supernovae that eject mass (e.g. Tan et al. 2001).

Acknowledgements. We thank Martin Laming for providing further results from, and information about, his modelling. The results

presented are based on observations obtained with XMM-Newton, an ESA science mission with instruments and contributions directly funded by ESA Member States and the USA.

References

- Anderson, M. C., & Rudnick, L. 1995, *ApJ*, 441, 307
 Aschenbach, B., Egger, R., & Trümper, J. 1995, *Nature*, 373, 587
 Bleeker, J. A. M., Willingale, R., van der Heyden, K. J., et al. 2001, *A&A*, 365, L225
 Braun, R. 1987, *A&A*, 171, 233
 Braun, R., Gull, S. F., & Perley, R. A. 1987, *Nature*, 327, 395
 Esteban, C., Smith, L. J., Vílchez, J. M., & Clegg, R. E. S. 1993, *A&A*, 272, 299
 Fabian, A. C., Willingale, R., Pye, J. P., Murray, S. S., & Fabbiano, G. 1980, *MNRAS*, 193, 175
 Fesen, R. A., Becker, R. H., & Goodrich, R. W. 1988, *ApJ*, 329, L89
 Hester, J. J., Scowen, P. A., Sankrit, R., et al. 1995, *ApJ*, 448, 240
 Hughes, J. P., Rakowski, C. E., Burrows, D. N., & Slane, P. O. 2000, *ApJ*, 528, L109
 Laming, J. M. 2001, *ApJ*, 563, 828
 Longair, M. S. 1994, *High Energy Astrophysics*, vol. 2 (Cambridge Univ. Press)
 Markert, T. H., Canizares, C. R., Clark, G. W., & Winkler, P. F. 1983, *ApJ*, 268, 13
 McKee, C. F. 1974, *ApJ*, 188, 335
 Reed, J. E., Hester, J. J., Fabian, A. C., & Winkler, P. F. 1995, *ApJ*, 440, 706
 Rosenberg, I. 1970, *MNRAS*, 151, 109
 Tan, J. C., Matzner, C. D., & McKee, C. F. 2001, *ApJ*, 551, 946-972
 Truelove, K. J., & McKee, C. F. 1999, *ApJ*, 120, 299
 Vink, J., Kaastra, J. S., & Bleeker, J. A. M. 1996, *A&A*, 307, L41
 Wang, L., & Wheeler, J. C. 1998, *ApJ*, 504, L87
 Willingale, R., Bleeker, J. A. M., van der Heyden, K. J., Kaastra, J. S., & Vink, J. 2002, *A&A*, 381, 1039