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THE SOFT X-RAY PROPERTIES OF A COMPLETE SAMPLE OF OPTICALLY SELECTED QUASARS. II. FINAL RESULTS

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ABSTRACT

We present the final results of a *ROSAT* PSPC program to study the soft X-ray emission properties of a complete sample of low- z quasars. This sample includes all 23 quasars from the Bright Quasar Survey with $z \leq 0.400$ and $N_{\text{HI}}^{\text{Gal}} < 1.9 \times 10^{20} \text{ cm}^{-2}$. Pointed *ROSAT* PSPC observations were made for all quasars, yielding high signal-to-noise (S/N) spectra for most objects, which allowed an accurate determination of the spectral shape. The following main results were obtained:

1. The spectra of 22 of the 23 quasars are consistent, to within $\sim 30\%$, with a single power-law model at rest-frame 0.2–2 keV. There is no evidence for significant soft excess emission with respect to the best-fit power law. We place a limit (95% confidence) of $\sim 5 \times 10^{19} \text{ cm}^{-2}$ on the amount of excess foreground absorption by cold gas for most of our quasars. The limits are $\sim 1 \times 10^{19} \text{ cm}^{-2}$ in the two highest S/N spectra.

2. The mean 0.2–2 keV continuum of quasars agrees remarkably well with an extrapolation of the mean 1050–350 Å continuum recently determined by Zheng et al. (1996) for $z > 0.33$ quasars. This suggests that there is no steep soft component below 0.2 keV.

3. Significant X-ray absorption ($\tau > 0.3$) by partially ionized gas (“warm absorber”) in quasars is rather rare, occurring for $\lesssim 5\%$ of the population, which is in sharp contrast to lower luminosity active galactic nuclei (AGNs), where significant absorption probably occurs for $\sim 50\%$ of the population.

4. Extensive correlation analysis of the X-ray continuum emission parameters with optical emission-line parameters indicates that the strongest correlation is between the spectral slope α_x and the $\text{H}\beta$ FWHM. A possible explanation for this remarkably strong correlation is a dependence of α_x on L/L_{Edd} , as seen in Galactic black hole candidates.

5. The strong correlations between α_x and $L_{[\text{O III}]}$, $\text{Fe II}/\text{H}\beta$, and the peak $[\text{O III}]$ to $\text{H}\beta$ flux ratio are verified. The physical origin of these correlations is still not understood.

6. There appears to be a distinct class of “X-ray-weak” quasars, which form $\sim 10\%$ of the population (three out of 23), where the X-ray emission is smaller, by a factor of 10–30, than expected based on their luminosity at other bands and on their $\text{H}\beta$ luminosity. These may be quasars in which the direct X-ray source is obscured and only scattered X-rays are observed.

7. Thin accretion disk models cannot reproduce the observed 0.2–2 keV spectral shape, and they also cannot reproduce the tight correlation between the optical and soft X-ray emission. An as yet unknown physical mechanism must be maintaining a strong correlation between the optical and soft X-ray emission.

8. The $\text{H I}/\text{He I}$ ratio in the high Galactic latitude ISM must be within 20%, and possibly within 5%, of the total H/He ratio of 10, which indicates that He in the diffuse H II gas component of the interstellar medium is mostly ionized to He II or He III.

We finally note the intriguing possibility that although $\langle \alpha_x \rangle$ in radio-loud quasars (-1.15 ± 0.14) is significantly flatter than in radio-quiet quasars (-1.72 ± 0.09) the X-ray emission may not be related to the presence of radio emission. The difference in $\langle \alpha_x \rangle$ may result from the strong α_x versus $\text{H}\beta$ FWHM correlation and the tendency of radio-loud quasars to have broader $\text{H}\beta$.

Subject headings: galaxies: active — galaxies: nuclei — quasars: general — X-rays: galaxies

1. INTRODUCTION

Quasars emit most of their power in the ultraviolet to soft X-ray regime. The position-sensitive proportional counter (PSPC) detector aboard *ROSAT* allowed a significantly improved study of the soft X-ray emission of quasars com-

pared with earlier missions (some of which were not sensitive below 2 keV), such as *HEAO-1*, *Einstein*, *EXOSAT*, and *Ginga* (e.g., Mushotzky 1984; Wilkes & Elvis 1987; Canizares & White 1989; Comastri et al. 1992; Lawson et al. 1992; Williams et al. 1992; and a recent review by

Mushotzky, Done, & Pounds 1993). These earlier studies indicated that the X-ray emission above 1-2 keV is well described by a power law with a spectral slope $\alpha_x = d \ln f_\nu / d \ln \nu$ of about -0.5 for radio-loud quasars and about -1.0 for radio-quiet quasars. Large heterogeneous samples of AGNs were recently studied using the *ROSAT* PSPC by Walter & Fink (1993) and by Wang, Brinkmann, & Bergeron (1996). However, the objects studied in the papers mentioned above do not form a complete sample, and the available results may be biased by various selection effects that were not well defined a priori. In particular, most of the studied objects are nearby, intrinsically X-ray-bright AGNs.

To overcome the potential biases in existing studies, we initiated a *ROSAT* PSPC program to make an accurate determination of the soft X-ray properties of a well-defined and complete sample of quasars, selected independently of their X-ray properties. This program was designed to address the following questions:

1. What are the soft X-ray spectral properties of the low-redshift quasar population?
2. Are simple thin accretion disk models (e.g., Laor 1990) able to fit the observed optical/UV/soft X-ray continuum? Are other modifying mechanisms, such as a hot corona (e.g., Czerny & Elvis 1987), required? Are models invoking optically thin free-free emission possible (e.g., Ferland, Korista, & Peterson 1990; Barvainis 1993)?
3. Do the observed soft X-ray properties display any significant correlations with other properties of these quasars? Are these correlations compatible with various models for the continuum and line emission mechanisms?

Our sample includes all 23 quasars from the Bright Quasar Survey (BQS) sample (Schmidt & Green 1983) with

$z \leq 0.400$ and $N_{\text{HI}}^{\text{Gal}} < 1.9 \times 10^{20} \text{ cm}^{-2}$, where $N_{\text{HI}}^{\text{Gal}}$ is the H I Galactic column density as measured at 21 cm. These selection criteria allow optimal study of soft X-ray emission at the lowest possible energy. The additional advantages of the BQS sample are that it has been extensively explored at other wavelengths (see Laor et al. 1994, hereafter Paper I, for further details) and that it includes only bright quasars, thus allowing high signal-to-noise (S/N) X-ray spectra for most objects. The sample selection criteria are independent of the quasar's X-ray properties, and we thus expect our sample to be representative of the low-redshift, optically selected quasar population.

Preliminary results from the analysis of the first 10 quasars available to us were described in Paper I. Here we report the analysis of the complete sample, which allows us to address the three questions posed above. The outline of the paper is as follows. In § 2 we describe the observations and the analysis of the spectra. Section 3 describes the analysis of correlations between the soft X-ray properties and other continuum and emission-line properties. In § 4 we compare our results with other soft X-ray observations and discuss some of the implications. We conclude in § 5 with answers to the questions raised above, and with some new questions to be addressed in future studies.

2. OBSERVATIONS AND ANALYSIS OF THE SPECTRA

The complete sample of 23 quasars is listed in Table 1 together with their redshifts, m_B and M_B magnitudes (calculated for $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $q = 0.5$), R (the radio-to-optical flux ratio), and $N_{\text{HI}}^{\text{Gal}}$. The redshifts and m_B and M_B magnitudes are taken from Schmidt & Green (1983); R is taken from Kellermann et al. (1989). Note that four of the 23 quasars in our sample are radio loud (defined here as $R \geq 10$).

TABLE 1
THE COMPLETE PG QUASAR SOFT X-RAY SURVEY SAMPLE^a

Quasar	Other Name	z^b	m_B^b	M_B^b	R^c	$N_{\text{HI}}^{\text{Gal}}$ ($\times 10^{20} \text{ cm}^{-2}$)	Reference ^d
PG 0947+396.....	K347-45	0.206	16.40	-24.06	0.25	1.92	1
PG 0953+414.....	K348-7	0.239	15.05	-25.73	0.44	1.12	2
PG 1001+054.....	...	0.161	16.13	-23.79	0.50	1.88	3
PG 1048+342.....	...	0.167	15.81	-24.19	$\lesssim 0.10$	1.74	1
PG 1114+445.....	...	0.144	16.05	-23.63	0.13	1.94	1
PG 1115+407.....	...	0.154	16.02	-23.81	0.17	1.74	1
PG 1116+215.....	TON 1388	0.177	15.17	-24.96	0.72	1.44	3
PG 1202+281.....	GQ COM	0.165	15.02	-24.96	0.19	1.72	3
PG 1216+069.....	...	0.334	15.68	-25.82	1.65	1.57	2
PG 1226+023.....	3C 273	0.158	12.86	-27.02	1138	1.68	4
PG 1309+355.....	TON 1565	0.184	15.45	-24.76	18.0	1.01	1
PG 1322+659.....	...	0.168	15.86	-24.15	0.12	1.88	3
PG 1352+183.....	PB 4142	0.158	15.71	-24.17	0.11	1.84	3
PG 1402+261.....	TON 182	0.164	15.57	-24.39	0.23	1.42	3
PG 1411+442.....	PB 1732	0.089	14.99	-23.65	0.13	1.05	2
PG 1415+451.....	...	0.114	15.74	-23.43	0.17	0.96	1
PG 1425+267.....	TON 202	0.366	15.67	-26.03	53.6	1.54	3
PG 1427+480.....	...	0.221	16.33	-24.28	0.16	1.69	1
PG 1440+356.....	Mrk 478	0.077	15.00	-23.32	0.37	0.97	1
PG 1444+407.....	...	0.267	15.95	-25.07	$\lesssim 0.08$	1.09	5
PG 1512+370.....	4C 37.43	0.371	15.97	-25.76	190	1.40	4
PG 1543+489.....	...	0.400	16.05	-25.84	0.15	1.68	2
PG 1626+554.....	...	0.133	16.17	-23.34	0.11	1.55	3

^a All PG sample quasars (i.e., $M_B < -23$) with $z \leq 0.400$ and $N_{\text{HI}}^{\text{Gal}} < 1.9 \times 10^{20}$.

^b From Schmidt & Green 1983; M_B calculated assuming $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $q_0 = 0.5$.

^c Radio (6 cm) to optical (4400 Å) flux ratio, taken from Kellermann et al. 1989. Radio-loud quasars have $R \geq 10$.

^d $N_{\text{HI}}^{\text{Gal}}$ taken from (1) Murphy et al. 1996; (2) Lockman & Savage 1995; (3) Elvis, Lockman, & Wilkes 1989; (4) Savage et al. 1993; (5) F. J. Lockman 1995, private communication.

The Galaxy becomes optically thick below 0.2 keV for the typical high Galactic column of $3 \times 10^{20} \text{ cm}^{-2}$ (Dickey & Lockman 1990; Morrison & McCammon 1983), and accurate values of $N_{\text{HI}}^{\text{Gal}}$ are therefore crucial even for our low- $N_{\text{HI}}^{\text{Gal}}$ sample. The $N_{\text{HI}}^{\text{Gal}}$ values given in column (7) of Table 1 are taken from Elvis, Lockman, & Wilkes (1989), Savage et al. (1993), Lockman & Savage (1995), and the recent extensive measurements by Murphy et al. (1996). All these measurements of $N_{\text{HI}}^{\text{Gal}}$ were made with the 140 foot (43 m) telescope of the NRAO at Green Bank, using the “bootstrapping” stray radiation correction method described by Lockman, Jahoda, & McCammon (1986), which provides an angular resolution of $21'$ and an uncertainty of $\Delta N_{\text{HI}}^{\text{Gal}} = 1 \times 10^{19} \text{ cm}^{-2}$ (and possibly lower for our low N_{H} quasars). This uncertainty introduces a flux error of 10% at 0.2 keV, 30% at 0.15 keV, and nearly a factor of 2 at 0.1 keV. Thus, reasonably accurate fluxes can be obtained down to ~ 0.15 keV. Note that Murphy et al. (1996) include accurate $N_{\text{HI}}^{\text{Gal}}$ measurements for about 220 AGNs, including most AGNs observed by *ROSAT*, which would be very useful for eliminating the significant systematic uncertainty in the PSPC spectral slope that must be present when a low-accuracy $N_{\text{HI}}^{\text{Gal}}$ is used.

Table 2 lists the PSPC observations of all the quasars. For the sake of completeness we include also the 10 quasars already reported in Paper I. All sources were detected, and their net source counts range from 93 to 38,015, with a median value of about 1900 counts. The PROS software package was used to extract the source counts. Table 2 includes the exposure times, the dates of the observations, the net number of counts and their statistical error, the count rate, the radius of the circular aperture used to extract the source counts, the offset of the X-ray position from the center of the PSPC field of view, the *ROSAT* sequence identification number, and the SASS version used

for the calibration of the data. All objects, except one, are typically within $15''$ of the center of the PSPC field of view, so all the identifications are secure. The one exception is PG 1440+356, where the pointing was offset by $40'$ from the position of the quasar (Gondhalekar et al. 1994). Note that the exposure times are uncertain by about 4% as the result of a number of possible systematic errors, as described by Fiore et al. (1994). The typically large number of counts for each object allows an accurate determination of the spectral slope for most objects, as described below.

Model fits to the extracted number of source counts per pulse-invariant (PI) channel, $N_{\text{ch}}^{\text{ob}}$, were carried out using the XSPEC software package. PI channels 1–12 of the original 256 channel spectra ($E < 0.11$ keV), were ignored, since they are not well calibrated and are inherently uncertain because of the large Galactic optical depth. The 1993 January PSPC calibration matrix was used for observations made after 1991 October 14, and earlier observations were fitted with the 1992 March calibration matrix. The best-fit model parameters are obtained by χ^2 minimization. Nearby channels were merged when $N_{\text{ch}}^{\text{ob}} < 10$. A 1% error was added in quadrature to the statistical error in $N_{\text{ch}}^{\text{ob}}$, to take into account possible systematic calibration errors (see Paper I for more complete details).

2.1. A Single Power Law

As in Paper I, we fitted each spectrum with a single power law of the form $f_E = e^{-N_{\text{H}}\sigma_E f_0} E^{\alpha_x}$, where f_E is the flux density, σ_E is the absorption cross section per H atom (Morrison & McCammon 1983), f_0 is the flux density at 1 keV, and E is in units of keV. We make three different fits for each object, with (1) N_{H} a free parameter, (2) $N_{\text{H}} = N_{\text{HI}}^{\text{Gal}}$, (3) $N_{\text{H}} = N_{\text{HI}}^{\text{Gal}}$, and $0.47 \leq E \leq 2.5$ keV, i.e., using only the hard *ROSAT* band (channels 12–34 of the rebinned 34 channel spectra). A comparison of fits 1 and 2 allows us to determine whether

TABLE 2
JOURNAL OF *ROSAT* PSPC OBSERVATIONS

Quasar (PG)	z	Exposure ^a (s)	Dates (Start–End)	Counts ^b	Counts s ⁻¹ ^b	r^c (arcmin)	Offset ^d (arcsec)	Sequence Number	SASS Version
0947+396.....	0.206	14283	1993 Oct 26–27	5345 ± 78	0.374	3.0	9.9	70136700	6.8
0953+414.....	0.239	7070	1992 Apr 30	4770 ± 74	0.675	3.0	26.0	700526	5.7
1001+054.....	0.161	8868	1992 May 17–Jun 2	93.4 ± 16	0.011	2.0	10.8	700774	5.7
1048+342.....	0.167	8072	1993 Nov 6	1624 ± 44	0.201	2.5	6.9	70136800	6.8
1114+445.....	0.144	7373	1993 Jun 6–7	777 ± 30	0.105	2.0	15.0	700800	6.4
1115+407.....	0.154	6495	1993 May 20	1934 ± 156	0.298	2.5	14.8	700801	6.4
1116+215.....	0.177	24942	1991 May 29–30	24272 ± 162	0.973	3.0	20.0	700228	5.4
1202+281.....	0.165	19777	1991 Jun 4–5	7882 ± 95	0.399	2.5	18.3	700232	5.3
1216+069.....	0.334	3464	1991 Dec 20–21	1370 ± 39	0.395	3.0	17.9	700021	5.6
1322+659.....	0.168	8389	1992 Nov 30–Dec 1	4187 ± 67	0.499	2.5	5.4	700803	6.2
1352+183.....	0.158	5521	1992 Jul 22	2285 ± 51	0.414	3.0	13.1	700804	5.8
1402+261.....	0.164	4357	1992 Jan 7–Jul 18	4153 ± 67	0.953	3.0	7.8	700226	5.11
1411+442.....	0.089	25319	1991 Jun 26–27	1031 ± 40	0.041	2.0	12.6	700248	5.3
1415+451.....	0.114	7432	1992 Jul 16	2514 ± 54	0.338	3.0	11.6	700805	5.8
1427+480.....	0.221	6780	1993 Jul 15–16	1104 ± 35	0.163	2.5	12.0	70137000	6.4
1440+356.....	0.077	2393	1992 Jan 17	4934 ± 74	2.062	7.0	24.0	700559	5.4
1444+407.....	0.267	5620	1993 Jul 19–20	1881 ± 46	0.335	2.5	12.0	70137100	6.4
1543+489.....	0.400	12759	1992 Aug 21–23	1375 ± 313	0.108	2.5	22.8	700808/9	5.8
1626+554.....	0.133	2312	1993 Aug 3–4	2294 ± 48	0.992	2.0	18.4	70137200	6.4
1226+023.....	0.158	6243	1991 Dec 14–15	38015 ± 197	6.089	3.0	8.0	700191	5.4
1309+355.....	0.184	2528	1992 Jul 6	593 ± 25	0.235	2.0	16.0	700802	5.7
1425+267.....	0.366	6716	1993 Jul 18–21	224 ± 18	0.033	1.5	18.0	70136900	6.4
1512+370.....	0.371	5230	1992 Aug 12–17	1498 ± 40	0.286	2.0	13.0	700807	5.8

^a Uncertain by about 4% due to systematic errors.

^b Net source counts in PI channels 11–245.

^c Effective aperture for integration of counts.

^d Offset of X-ray position from the center of the PSPC field of view.

there is evidence for a significant intrinsic absorption excess or emission excess relative to a single power-law fit with $N_{\text{H}} = N_{\text{HI}}^{\text{Gal}}$. This comparison also allows us, as is further shown in § 4.2, to determine whether the 21 cm measurement of $N_{\text{HI}}^{\text{Gal}}$ is a reliable measure of the Galactic soft X-ray opacity. A comparison of fits 2 and 3 allows us to look for a dependence of the power-law slope on energy.

Table 3 provides the results of fits 1-3 described above. The table includes the 13 objects not reported in Paper I, and six objects from Paper I for which we only now have the accurate $N_{\text{HI}}^{\text{Gal}}$ values. For each fit we give the best-fitting spectral slope α_x , the normalization of the power-law flux at 1 keV (f_0), the best-fitting N_{H} , the χ^2 of the fit (χ_{fit}^2), the number of degrees of freedom (dof), and the probability for

TABLE 3
SINGLE POWER-LAW MODELS ($f_E = e^{-N_{\text{H}}\sigma_E} f_0 E^{\alpha_x}$)

Name (PG)	Fit ^a	α_x^b	f_0^c ($\times 10^{-4}$)	N_{H}^d (10^{20} cm^{-2})	$\chi_{\text{fit}}^2/\text{dof}$	$\text{Pr}(\chi^2 \geq \chi_{\text{fit}}^2)$
0947+396.....	1	$-1.36_{-0.09}^{+0.10}$	7.64 ± 0.27	1.54 ± 0.22	26.51/26	0.43
	2	-1.51 ± 0.02	7.64 ± 0.20	1.92	32.10/27	0.23
	3	-1.44 ± 0.07	7.68 ± 0.18	1.92	21.27/18	0.27
0953+414.....	1	-1.71 ± 0.12	9.30 ± 0.50	1.50 ± 0.25	14.31/26	0.97
	2	-1.57 ± 0.03	9.33 ± 0.14	1.12	18.12/27	0.90
	3	-1.62 ± 0.11	9.25 ± 0.32	1.12	9.56/18	0.95
1001+054.....	1	$-2.80_{-0.58}^{+1.07}$	$0.048_{-0.039}^{+0.06}$	1.88	2.02/4	0.73
1048+342.....	1	-1.27 ± 0.21	4.23 ± 0.32	1.47 ± 0.45	14.86/23	0.90
	2	-1.39 ± 0.05	4.24 ± 0.21	1.74	15.68/24	0.90
	3	-1.33 ± 0.15	4.24 ± 0.20	1.74	10.84/14	0.70
1114+445.....	1	$-1.13_{-0.31}^{+0.24}$	1.97 ± 0.20	1.04 ± 0.62	59.72/25	1.1×10^{-4}
	2	-1.55 ± 0.06	1.91 ± 0.13	1.94	63.85/26	5×10^{-5}
	3	-0.17 ± 0.31	1.96 ± 0.13	1.94	36.16/17	4.4×10^{-3}
1115+407.....	1	-1.77 ± 0.17	3.82 ± 0.28	1.48 ± 0.36	24.54/21	0.27
	2	-1.89 ± 0.04	3.79 ± 0.19	1.74	25.73/22	0.26
	3	-1.86 ± 0.14	3.81 ± 0.20	1.74	11.28/13	0.59
1116+215.....	1	-1.73 ± 0.05	13.1 ± 0.28	1.43 ± 0.10	38.13/29	0.12
	2	-1.73 ± 0.01	13.1 ± 0.16	1.44	38.16/30	0.15
	3	-1.69 ± 0.04	13.2 ± 0.15	1.44	24.04/21	0.29
1202+281.....	1	-1.23 ± 0.08	9.93 ± 0.27	1.88 ± 0.20	20.79/28	0.83
	2	-1.22 ± 0.02	9.87 ± 0.17	1.72	22.45/29	0.80
	3	-1.22 ± 0.05	9.88 ± 0.17	1.72	18.83/20	0.53
1216+069.....	1	-1.52 ± 0.20	8.01 ± 0.57	$1.96_{-0.43}^{+0.48}$	24.14/25	0.51
	2	-1.36 ± 0.04	7.93 ± 0.38	1.57	26.01/26	0.46
	3	-1.43 ± 0.13	7.93 ± 0.41	1.57	17.05/17	0.45
1309+335.....	1	-1.725 ± 0.31	3.00 ± 0.40	$1.43_{-0.50}^{+0.62}$	20.99/22	0.52
	2	-1.51 ± 0.07	3.02 ± 0.26	1.01	22.63/23	0.48
	3	-1.42 ± 0.25	3.03 ± 0.40	1.01	10.92/14	0.69
1402+261.....	1	-1.79 ± 0.13	10.00 ± 0.67	1.16 ± 0.22	26.02/25	0.41
	2	-1.93 ± 0.03	9.86 ± 0.04	1.42	29.13/26	0.31
	3	-1.57 ± 0.14	10.3 ± 0.4	1.42	10.16/16	0.86
1411+442.....	1	$-2.14_{-0.29}^{+0.27}$	0.31 ± 0.05	1.42 ± 0.54	25.64/20	0.18
	2	-1.97 ± 0.07	0.31 ± 0.03	1.05	26.46/21	0.19
	3	-1.87 ± 0.27	0.30 ± 0.03	1.05	14.53/12	0.27
1415+451.....	1	-1.59 ± 0.13	3.32 ± 0.25	$0.67_{-0.19}^{+0.22}$	19.98/24	0.70
	2	-1.74 ± 0.04	3.28 ± 0.18	0.96	22.91/25	0.58
	3	-1.57 ± 0.14	3.40 ± 0.17	0.96	15.31/16	0.50
1425+267.....	1	$-0.71_{-0.58}^{+0.85}$	0.84 ± 0.18	$1.0_{-1.0}^{+2.0}$	13.2/9	0.15
	2	-0.94 ± 0.16	0.85 ± 0.11	1.54	13.47/10	0.20
	3	-0.13 ± 0.64	0.81 ± 0.13	1.54	5.50/4	0.24
1427+480.....	1	-1.37 ± 0.27	3.25 ± 0.31	1.60 ± 0.60	19.82/23	0.65
	2	-1.41 ± 0.05	3.26 ± 0.22	1.69	19.98/24	0.70
	3	-1.24 ± 0.20	3.28 ± 0.20	1.69	13.79/15	0.54
1440+356.....	1	-2.38 ± 0.13	19.2 ± 1.5	$1.49_{-0.20}^{+0.22}$	20.53/24	0.66
	2	-2.08 ± 0.03	20.2 ± 0.9	0.97	37.60/25	0.05
	3	-2.17 ± 0.13	19.3 ± 1.0	0.97	11.21/16	0.80
1444+407.....	1	-1.86 ± 0.20	2.91 ± 0.34	$1.00_{-0.29}^{+0.35}$	14.46/19	0.76
	2	-1.91 ± 0.05	2.89 ± 0.22	1.09	14.66/20	0.80
	3	-2.14 ± 0.19	2.82 ± 0.23	1.09	4.39/11	0.96
1543+489.....	1	$-2.57_{-0.31}^{+0.27}$	0.98 ± 0.13	$2.67_{-0.56}^{+0.67}$	29.93/23	0.15
	2	-2.11 ± 0.05	1.02 ± 0.08	1.68	38.62/24	0.03
	3	-2.19 ± 0.20	0.99 ± 0.10	1.68	20.95/15	0.14
1626+554.....	1	-1.68 ± 0.14	12.9 ± 0.9	1.26 ± 0.29	17.78/23	0.77
	2	-1.94 ± 0.04	12.8 ± 0.6	1.82	25.26/24	0.39
	3	-1.83 ± 0.15	13.1 ± 0.6	1.82	16.49/15	0.35

^a The best-fit single power-law model for (1) PI channels 11–245, and N_{H} a free parameter; (2) PI channels 11–245, and $N_{\text{H}} = N_{\text{HI}}^{\text{Gal}}$; (3) PI channels 47–245, and $N_{\text{H}} = N_{\text{HI}}^{\text{Gal}}$.

^b The error is 1σ for either two interesting parameters (fit 1) or one interesting parameter (fits 2 and 3).

^c In units of $\text{keV cm}^{-2} \text{ s}^{-1} \text{ keV}^{-1}$, or, equivalently, $6.626 \times 10^{-27} \text{ ergs cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1}$, at 1 keV.

^d In fit 1 the N_{H} error is 1σ for two interesting parameters. In fits 2 and 3 the N_{H} error is assumed to be zero.

$\chi^2 \geq \chi_{\text{fit}}^2$. The errors $\Delta\alpha_x$ and ΔN_{H} in fit 1 were calculated by making a grid search for models with $\Delta\chi^2 = 2.30$, as appropriate for a 1σ confidence level for two interesting parameters (e.g., Press et al. 1989). The error on the slope, $\Delta\alpha_x$, in fits 2 and 3 is calculated by requiring $\Delta\chi^2 = 1.0$ (i.e., 68% for one interesting parameter). We neglect the effect of $\Delta N_{\text{H}}^{\text{Gal}}$ on $\Delta\alpha_x$ in fits 2 and 3, since we use accurate values of $N_{\text{H}}^{\text{Gal}}$ for all objects.

The observed and best-fit spectra for the 13 quasars not reported in Paper I are displayed in Figure 1. There are three panels for each object. The upper panel displays the observed count rate per keV as a function of channel energy; the histogram represents the expected rate from the best-fit power-law model, $N_{\text{ch}}^{\text{ob}}$, with N_{H} a free parameter (fit 1). The middle panel displays Δ/σ , where $\Delta = N_{\text{ch}}^{\text{ob}} - N_{\text{ch}}^{\text{mod}}$ and σ is the standard error in $N_{\text{ch}}^{\text{ob}}$. This plot helps indicate what features in the spectrum are significant. The lower

panel displays the fractional deviations from the expected flux or, equivalently, $\Delta/N_{\text{ch}}^{\text{mod}}$, which indicates the fractional amplitude of the observed features.

As shown in Table 3, in all 13 objects the simple power-law model with $N_{\text{H}}^{\text{Gal}}$ (fit 2) provides an acceptable fit [i.e., $\text{Pr}(\chi^2 \geq \chi_{\text{fit}}^2) > 0.01$]. Note in particular the spectrum of PG 1116+215, which despite the very high S/N available (24,272 net counts) shows no deviations from a simple power law above a level of $\lesssim 10\%$. In Paper I a simple power-law model could not provide an acceptable fit to three of the 10 quasars, although in two of them the apparent features could not be fitted with a simple physical model, and in one of them this may be due to calibration errors (see § 4.1).

As mentioned above, a comparison of the free N_{H} fit (fit 1) with the $N_{\text{H}}^{\text{Gal}}$ fit (fit 2) allows us to look for evidence for an absorption or an emission excess. We measure the statistical

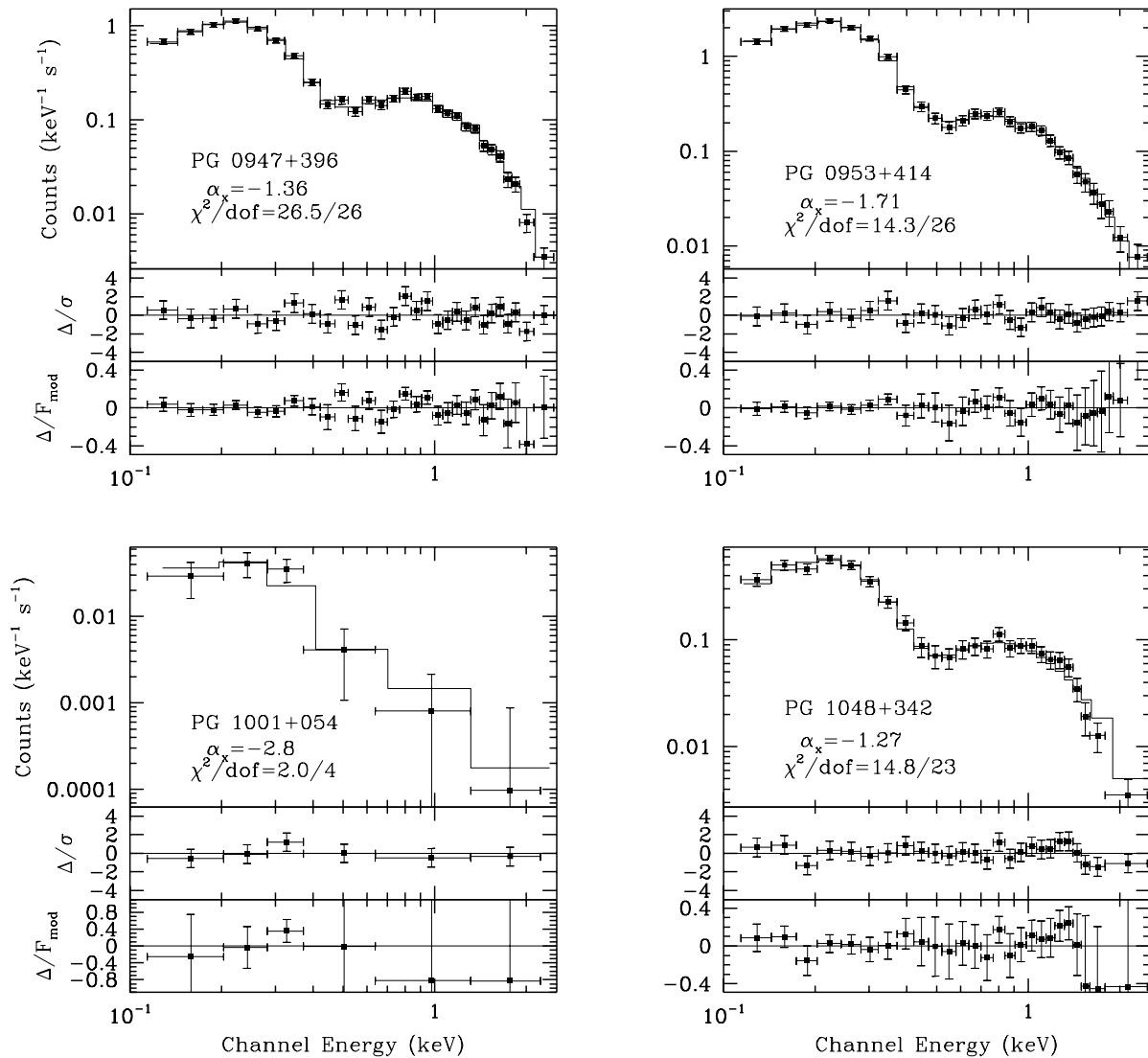


FIG. 1a

FIG. 1.—Observed vs. single power-law fit PSPC spectra of the 13 quasars not reported in Paper I. The upper panel for each object shows the observed count rate (points with error bars), and the expected count rate using the best-fit single power-law model with a free N_{H} (solid-line histogram). The middle panel shows the observed minus the expected flux in units of standard deviations, and the lower panel for each object shows the fractional deviations from the expected flux. The best-fit α_x , the fit χ^2 , and the number of degrees of freedom (dof) are indicated for each object. The spectra of all quasars are consistent with a simple power law. Note in particular the high S/N spectrum of PG 1116+215, where intrinsic features below 1 keV must have an amplitude of less than $\approx 10\%$.

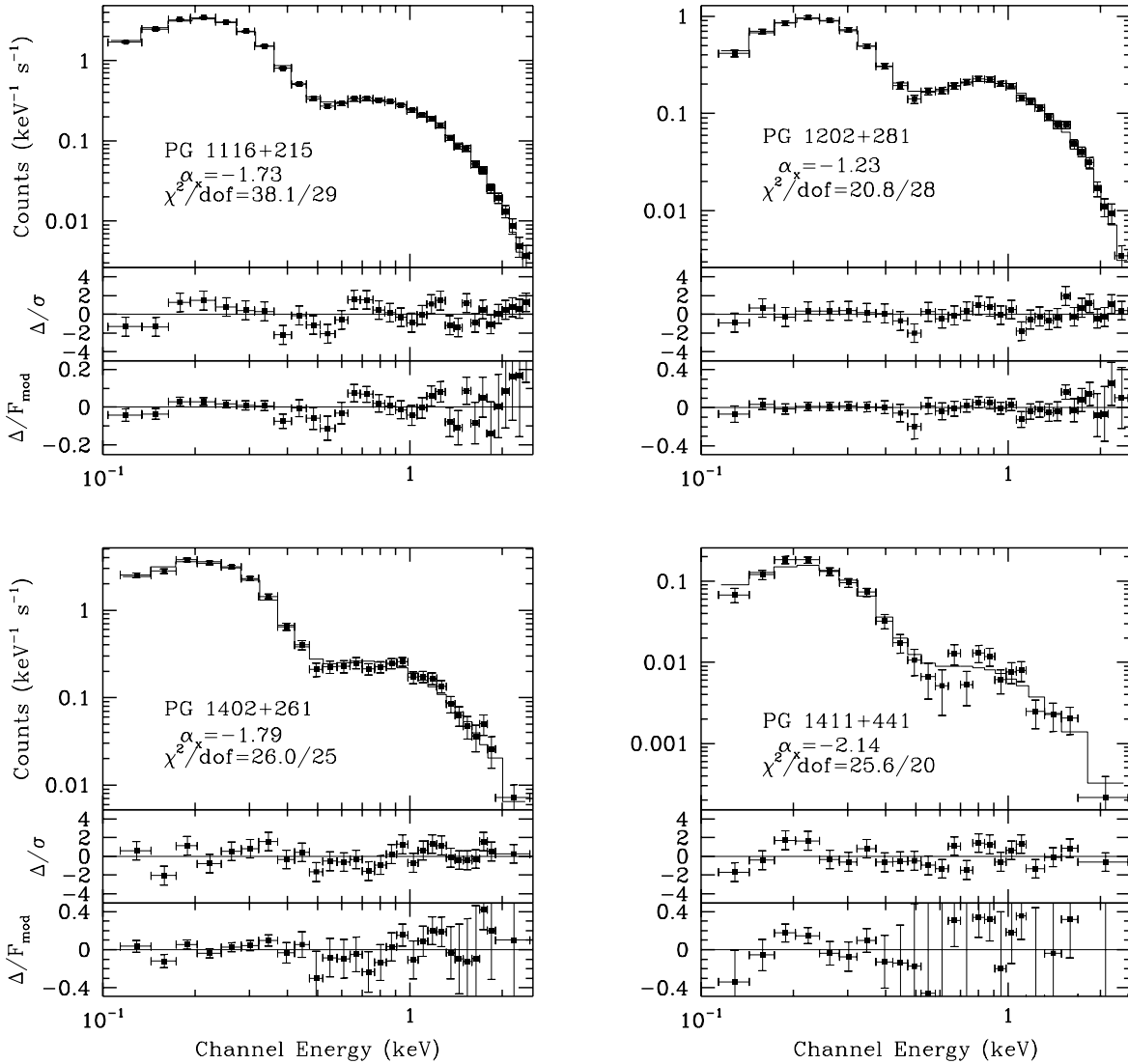


FIG. 1b

significance of the reduction in χ^2_{fit} with the addition of N_{H} as a free parameter using the F -test (Bevington 1969). In PG 1440+356 we find a significant reduction with $\text{Pr} = 7.5 \times 10^{-4}$ ($F = 12.98$ for 26 dof), where Pr is the probability that the reduction in χ^2 is not statistically significant (calculated using the FTEST routine in Press et al. 1989). The N_{H} obtained in fit 1 suggests intrinsic absorption of about $5 \times 10^{19} \text{ cm}^{-2}$ above the Galactic absorption. However, unlike all other objects, PG 1440+356 was observed significantly off axis (see Table 2), and some small systematic calibration errors may be present there. Note also that the χ^2_{fit} of the fit with $N_{\text{H}}^{\text{Gal}}$ ($\text{Pr} = 0.05$) is still acceptable. We therefore cannot conclude that an extra absorber must be present in PG 1440+356. Marshall et al. (1996) found a very steep slope ($\alpha = -4.7 \pm 0.65$) in PG 1440+35 at 0.1–0.15 keV (80–120 Å) using the *Extreme Ultraviolet Explorer*. There is no indication for such a component in the PSPC spectrum below 0.15 keV (Fig. 1c). However, given the very low sensitivity of the PSPC below 0.15 keV, such a steep soft component may still be consistent with the PSPC spectrum. In the other 12 objects the free N_{H} fit does not provide a significant improvement (i.e.,

$\text{Pr} > 0.01$), and thus there is no clear evidence for either intrinsic absorption or low-energy excess emission above a simple power law.

Figure 2 compares the Galactic N_{H} deduced from the accurate 21 cm measurements with the best-fit X-ray column deduced using the free N_{H} fit. The straight line represents equal columns. The χ^2 of the $N_{\text{H}}(21 \text{ cm}) = N_{\text{H}}(\text{X-ray})$ model is 31.9 for 22 dof (PG 1001+054 was not included because of the low S/N), which is acceptable at the 8% level. This result demonstrates that there is no significant excess absorption over the Galactic value in any of our objects. It is interesting to note that in our highest S/N spectra, those of PG 1116+215 and PG 1226+023, $N_{\text{H}}(\text{X-ray})$ is determined to a level of $(0.8\text{--}1) \times 10^{19} \text{ cm}^{-2}$ and is still consistent with $N_{\text{H}}(21 \text{ cm})$, indicating that the two methods agree to better than 10%.

The average hard *ROSAT* band (0.5–2 keV) slope for the complete sample is -1.59 ± 0.08 (excluding PG 1114+445, which is affected by a warm absorber, and PG 1001+054 and PG 1425+267, for which the S/N is very low). This slope is not significantly different from the average slope for the full *ROSAT* band, -1.63 ± 0.07 .

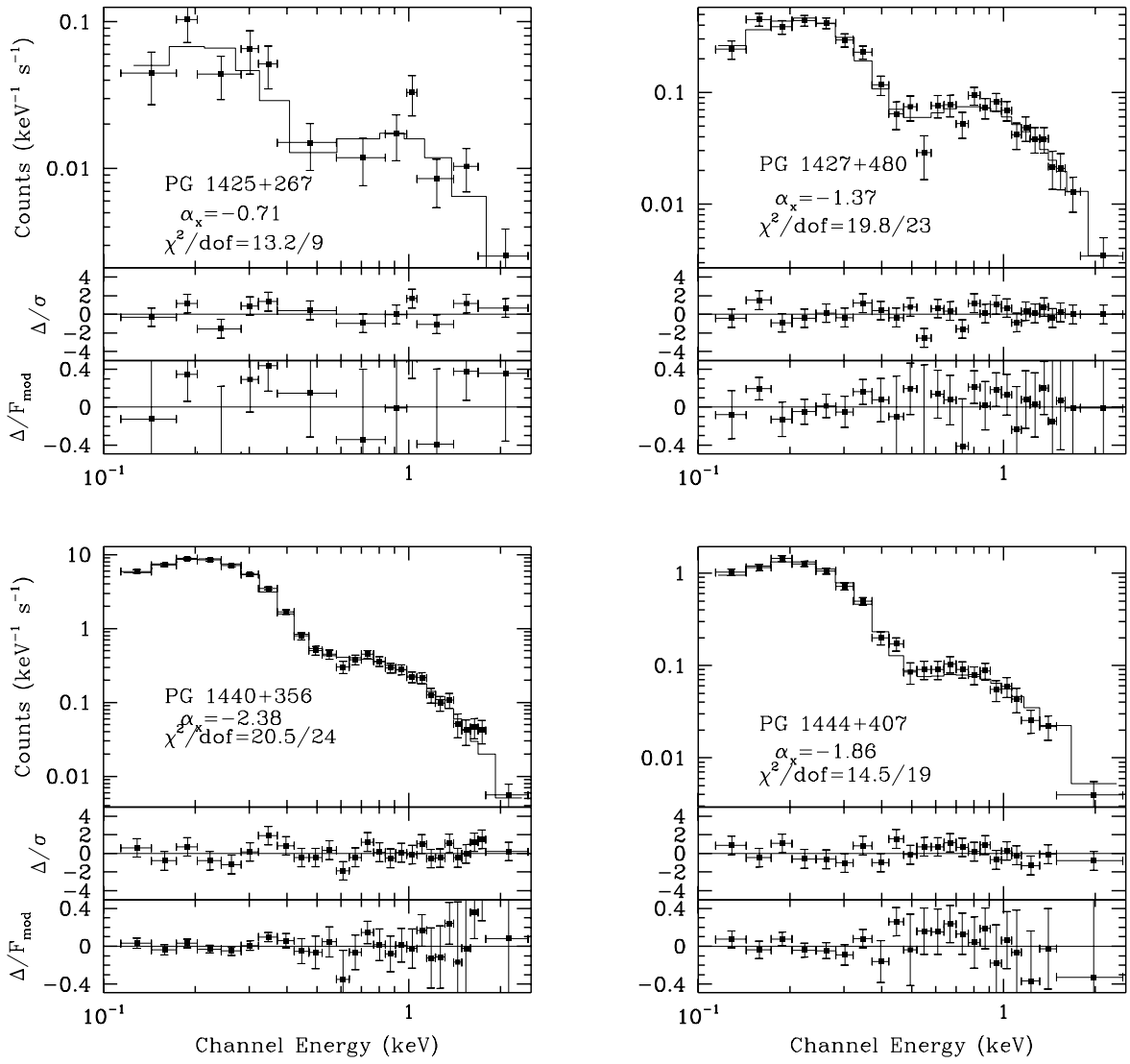


FIG. 1c

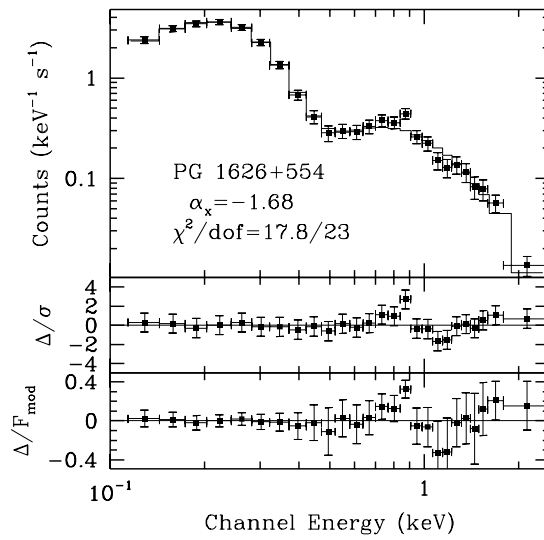


FIG. 1d

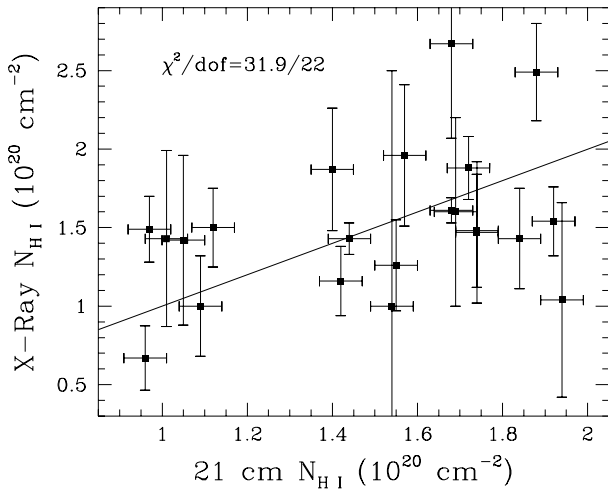


FIG. 2.—The H I column determined by 21 cm measurements vs. the best-fit H I column determined by a power-law fit to the quasars' PSPC spectrum. The assumption that both values are equal, indicated by the straight line, is acceptable at the 8% level. Note, in particular, the two highest S/N objects, which deviate from the straight line by less than $1 \times 10^{19} \text{ cm}^{-2}$. The agreement between the two measures of N_{H} indicates a lack of intrinsic cold gas absorption in quasars, and that $\text{H I}/\text{He I} \approx \text{H}/\text{He}$ in the ISM at high Galactic latitudes.

Spectral fits to the PSPC data of some of the objects in our complete sample have already been reported by Gondhalekar et al. (1994), Ulrich & Molendi (1996), Rachen, Mannheim, & Biermann (1996), and Wang et al. (1996). The results of the single power-law fit with a free N_{H} in these papers are all consistent with our results for the overlapping objects. The only discrepancy is with PG 1444+407 for which, although both the present authors and Wang et al. (1996) find a similar slope, Wang et al. find evidence for absorption while we find no such evidence. For a simple power-law fit to PG 1444+407 Wang et al. (1996) find $\chi^2 = 26$ for 20 dof, which is acceptable only at the 17% level ($\sim 1.4 \sigma$ level), while we find for such a fit $\chi^2 = 14.7$ for 20 dof, which is acceptable at the 80% level.

As discussed in Paper I (§ 5.1.3), the difference in spectral slopes at hard (2–10 keV) and soft X-rays raises the possibility that α_x may be changing within the PSPC band itself. The individual spectra are well fitted by a simple power law, and thus any spectral curvature must be consistent with zero. Stronger constraints on the spectral curvature may be obtained by measuring the average curvature parameter (β , defined in Paper I) for the complete sample, since the random error in the mean is smaller by $N^{1/2} \sim 5$ than the random error for individual objects. Unfortunately, the PSPC calibration uncertainty at low energy, discussed in Paper I, introduces a systematic error in β (which obviously does not cancel out as $N^{1/2}$) and, as shown in Paper I, does not allow a reliable determination of the curvature parameter. We therefore did not try to constrain the spectral curvature parameter in this paper.

3. CORRELATION ANALYSIS

Table 4 presents eight of the 12 rest-frame continuum parameters and seven of the 18 emission-line parameters used for the correlation analysis. The spectral slopes are defined [the flux subscript indicates $\log \nu(\text{Hz})$] as $\alpha_o = \log(f_{1.5}/f_{14.5})/0.5$ (1–0.3 μm), $\alpha_{ox} = \log(f_{17.685}/f_{15})/2.685$ (3000 \AA –2 keV), and $\alpha_{os} = \log(f_{16.861}/f_{15})/1.861$ (3000

\AA –0.3 keV). The X-ray continuum parameters are from fit 2 and from Paper I. The near-IR and optical continuum parameters are taken from Neugebauer et al. (1987). The emission-line parameters were taken from Boroson & Green (1992). Luminosities were calculated assuming $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $q_0 = 0.5$.

The four additional continuum parameters that are omitted from Table 5 for the sake of brevity are $\alpha_{irx} = \log(f_{16.861}/f_{14.25})/2.611$ (1.69 μm –2 keV), $\alpha_{irs} = \log(f_{17.685}/f_{14.25})/3.435$ (1.69 μm –0.3 keV), radio luminosity (Kellermann et al. 1989), and 1 μm luminosity (Neugebauer et al. 1987). The 11 emission-line parameters not detailed in Table 5 are [O III] EW, Fe II EW, H β EW, He II EW, [O III]/H β and He II/H β flux ratios, [O III] peak flux to H β peak flux ratio, radio to optical flux ratio, and the H β asymmetry, shape, and shift parameters. All these 11 parameters are listed in Table 2 of Boroson & Green (1992).

The significance of the correlations was tested using the Spearman rank-order correlation coefficient (r_s), which is sensitive to any monotonic relation between the two variables. A summary of the main correlation coefficients and their two-sided significance is given in Table 5.

3.1. Significance Level

In Paper I the correlation analysis was carried out using 10 objects, and only relatively strong correlations ($r_s \geq 0.76$) could be detected at the required significance level ($\text{Pr} \leq 0.01$). Here, with 23 objects, $\text{Pr} \leq 0.01$ corresponds to $r_s \geq 0.52$, and we can thus test for the presence of weaker correlations and check whether the correlations suggested in Paper I remain significant. We have searched for correlations among the 12 continuum emission parameters, and between these 12 parameters and the 18 emission-line parameters listed above, which gives a total of 294 different correlations. One thus expects about 1 spurious correlation with $\text{Pr} \leq 3.4 \times 10^{-3}$ in our analysis, and for a significance level of 1% one would now have to go to $\text{Pr} \leq 3.4 \times 10^{-5}$ rather than $\text{Pr} \leq 1 \times 10^{-2}$. However, we find that there are actually 42 correlations with $\text{Pr} \leq 3.4 \times 10^{-3}$, rather than just one, in our sample. Thus, the probability that any one of them is the spurious one is only 2.4%, and the significance level of these correlations is reduced by a factor of 7 ($= 0.024/0.0034$) rather than a factor of 300. Below we assume that correlations with $\text{Pr} \leq 1 \times 10^{-3}$ are significant at the 1% level (there are 30 correlations with $\text{Pr} \leq 1 \times 10^{-3}$, versus an expected number of 0.3). Thus, given the large number of correlations we looked at, we can only test reliably for correlations with $r_s \geq 0.64$ (which corresponds to $\text{Pr} \leq 1 \times 10^{-3}$ for 23 data points).

3.2. The Near-IR to X-Ray Energy Distribution

A comparison of the rest-frame spectral energy distributions of all 23 quasars is shown in Figure 3. The 3–0.3 μm continuum is from Neugebauer et al. (1987), and the 0.2–2 keV continuum is from Paper I and from this paper. The upper panel shows the absolute luminosities, and the two lower panels the luminosity normalized to unity at $\log \nu = 14.25$ for radio-quiet quasars and for radio-loud quasars. Note the relatively small dispersion in the normalized 0.3 keV ($\log \nu = 16.861$) luminosity. The outlying objects are labeled. PG 1626+554 is the only object where a steep α_x is clearly associated with a strong soft excess (relative to the near-IR flux). In other objects a steep α_x tends to be associated with a low 2 keV flux. This trend is

TABLE 4
CONTINUUM AND EMISSION-LINE PARAMETERS

Name (PG)	$A(z)^a$	$L_{1.69}$ μm^b	$L_{30.00}$ \AA^b	$L_{0.3\text{ keV}}$	$L_{2\text{ keV}}$	α_o	α_x	α_{ox}	α_{os}	$I_{\text{H}\beta}^{b,c}$	$I_{\text{Fe II}}^{b,c}$	$L_{\text{He II}}^{b,c}$	$L_{[\text{O III}]}^{b,c}$	FWHM ^d	H β	$R(\text{Fe II})^e$	$T_{\text{rot}} \leq$ 10^5 K
0947+396 ...	56.30	44.92	45.21	44.70	44.28	-0.211	-1.510	-1.348	-1.277	43.36	42.71	42.59	42.49	4830	4830	0.23	5.3
0953+414 ...	56.44	45.43	45.88	44.96	44.49	-0.119	-1.570	-1.518	-1.495	44.11	43.51	42.85	43.17	3130	3130	0.25	4.1
1001+054 ...	56.08	44.94	45.21	43.02	41.54	-0.239	-2.800	-2.369	-2.178	43.25	43.17	42.15	42.15	1740	1740	0.82	2.9
1048+342 ...	56.11	44.74	45.14	44.17	43.85	-0.351	-1.390	-1.481	-1.522	43.62	43.12	42.76	42.70	3600	3600	0.32	4.5
1114+445 ...	55.98	45.00	45.08	43.79	43.89	-0.574	-0.880	-1.443	-1.692	43.29	42.59	41.85	42.52	4570	4570	0.20	4.2
1115+407 ...	56.04	44.73	44.96	44.34	43.61	-0.505	-1.890	-1.502	-1.320	42.92	42.65	40.83 ^g	41.91	1720	1720	0.54	7.1
1116+215 ...	56.16	45.55	45.88	44.92	44.32	-0.146	-1.730	-1.581	-1.515	44.14	43.80	42.44	42.89	2920	2920	0.47	4.0
1202+281 ...	56.10	45.04	45.00	44.43	44.25	-0.680	-1.220	-1.280	-1.306	43.27	42.73	42.05	42.67	5050	5050	0.29	8.3
1216+069 ...	56.74	45.58	45.94	45.08	44.78	-0.380	-1.360	-1.430	-1.461	44.11	43.40	42.83	43.17	5190	5190	0.20	5.4
1226+023 ...	56.06	46.15	46.45	45.48	45.53	-0.281	-0.942	-1.343	-1.521	44.59	44.34	43.06	43.23	3520	3520	0.57	4.0
1309+335 ...	56.20	45.06	45.32	44.19	43.77	-0.585	-1.510	-1.580	-1.612	43.25	42.69	41.24 ^g	42.82	2940	2940	0.28	5.2
1322+659 ...	56.12	44.92	45.26 ^f	44.64	44.07	-0.231	-1.690	-1.441	-1.331	43.20	42.97	42.61	42.22	2790	2790	0.59	5.1
1352+183 ...	56.06	44.87	45.26	44.47	44.04	-0.262	-1.524	-1.456	-1.426	43.45	43.11	42.24	42.33	3600	3600	0.46	4.6
1402+261 ...	56.10	45.15	45.45	44.84	44.08	-0.108	-1.930	-1.513	-1.329	43.35	43.44	41.16 ^g	41.46	1910	1910	1.23	4.7
1411+442 ...	55.55	44.86	44.93	42.79	41.99	-0.613	-1.970	-2.095	-2.150	43.17	42.86	41.95	42.32	2670	2670	0.49	3.2
1415+451 ...	55.77	44.84	44.81	43.91	43.30	-0.893	-1.740	-1.561	-1.482	42.80	42.90	41.40	41.03	2620	2620	1.25	9.3
1425+267 ...	56.83	45.46	45.81	43.92	43.97	-0.478	-0.940	-1.684	-2.014	43.95	42.98	42.23	43.53	9410	9410	0.11	3.6
1427+480 ...	56.36	44.78	45.22	44.33	43.99	-0.152	-1.410	-1.458	-1.479	43.42	42.99	42.61	43.05	2540	2540	0.36	4.1
1440+356 ...	55.42	44.91	44.91	44.53	43.64	-0.724	-2.080	-1.473	-1.204	42.91	42.98	41.67	42.10	1450	1450	1.19	15.0
1444+407 ...	56.54	45.38	45.63	44.77	44.02	-0.278	-1.910	-1.598	-1.460	43.50	43.71	41.39 ^g	41.69	2480	2480	1.45	5.0
1512+370 ...	56.84	45.39	45.89	44.96	44.79	-0.068	-1.210	-1.412	-1.501	44.06	43.27 ^g	42.93	43.73	6810	6810	0.16 ^g	4.1
1543+489 ...	56.91	45.60	45.80	44.86	43.95	-0.192	-2.110	-1.691	-1.506	43.77	43.70	41.56 ^g	41.56 ^g	1560	1560	0.85	4.8
1626+554 ...	55.91	44.69	45.02	44.76	43.99	-0.330	-1.940	-1.385	-1.139	43.27	42.78	42.19	42.07	4490	4490	0.32	7.5
Average.....	...	45.13	45.39	44.43	43.92	-0.36	-1.62	-1.55	-1.52	43.51	43.14	42.36	42.51	3545	3545	0.57	5.5
Dispersion....	...	± 0.37	± 0.43	± 0.63	± 0.82	± 0.22	± 0.45	± 0.24	± 0.26	± 0.46	± 0.45	± 0.46	± 0.67	± 1865	± 1865	± 0.39	± 2.6

NOTE—All continuum luminosities are rest-frame νL_ν . Continuum and line luminosities are in units of $\log(\text{ergs s}^{-1})$. Luminosities were calculated assuming $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $q_0 = 0.5$.

^a $A(z) = \log(\nu L_\nu/\nu'_L) = \log \left[\frac{4\pi^2 c^2}{H_0^2} [1+z-(1+z)^{1/2}]^2 \right]$.

^b Calculated using continuum fluxes from Neugebauer et al. 1987.

^c Equivalent widths from Boroson & Green 1992.

^d In km s^{-1} .

^e Fe II/H β flux ratio.

^f D. Mào 1993 (private communication).

^g 2σ upper limit.

TABLE 5
SPEARMAN RANK-ORDER CORRELATION COEFFICIENTS (r_s) AND PROBABILITIES (Pr)

Parameter	Statistic	z	$L_{1.69 \mu\text{m}}$	$L_{3.000 \text{ \AA}}$	$L_{2 \text{ keV}}$	α_o	α_x	α_{ox}	α_{os}	$L_{\text{H}\beta}$	$L_{\text{Fe II}}$	$L_{\text{He II}}$	$L_{[\text{O III}]}$	H β (FWHM)	Fe II/H β
$L_{0.3 \text{ keV}}$...	r_s	0.443	0.594	0.711	0.837	0.567	0.074	0.378	0.375	0.654	0.615	0.714	0.283	0.123	0.076
	Pr	0.034	0.0028	1.4×10^{-4}	6.5×10^{-7}	0.0047	0.737	0.075	0.078	7.1×10^{-4}	0.002	8.6×10^{-4}	0.202	0.576	0.736
$L_{2 \text{ keV}}$...	r_s	0.488	0.576	0.725	1	0.564	0.482	0.585	0.279	0.734	0.444	0.769	0.543	0.514	-0.288
	Pr	0.018	0.004	9.2×10^{-5}	0	0.005	0.020	0.0034	0.199	6.6×10^{-5}	0.038	2×10^{-4}	0.009	0.012	0.194
α_x ...	r_s	0.302	0.262	0.347	0.482	-0.049	1	0.533	-0.201	0.458	-0.167	0.411	0.754	0.791	-0.714
	Pr	0.162	0.226	0.105	0.020	0.825	0	0.0089	0.358	0.028	0.455	0.090	5×10^{-5}	6.9×10^{-6}	1.9×10^{-4}
α_{ox} ...	r_s	-0.090	-0.095	0.109	0.585	0.010	0.533	1	0.556	0.117	-0.256	0.336	0.224	0.535	-0.337
	Pr	0.684	0.666	0.933	0.0034	0.964	0.0089	0	0.0058	0.595	0.249	0.173	0.315	0.0085	0.125
α_{os} ...	r_s	-0.148	-0.308	-0.249	0.279	0.045	-0.201	0.556	1	-0.244	-0.167	-0.036	-0.432	-0.079	0.167
	Pr	0.500	0.153	0.252	0.198	0.840	0.358	0.0058	0	0.261	0.458	0.888	0.045	0.719	0.458
α_{FeX} ...	r_s	0.137	-0.025	0.259	0.732	0.329	0.549	0.914	0.470	0.364	-0.045	0.631	0.409	0.557	-0.387
	Pr	0.533	0.910	0.233	7.2×10^{-5}	0.125	0.0066	1.1×10^{-9}	0.023	0.087	0.841	0.005	0.059	0.0057	0.075

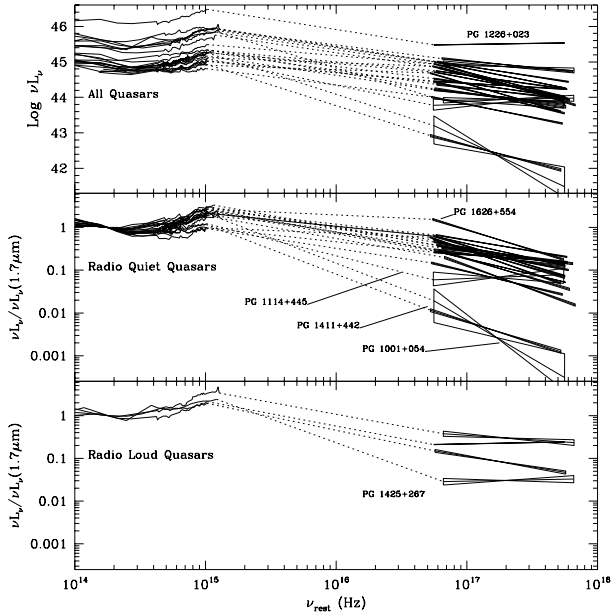


FIG. 3.—Comparison of the spectral energy distribution of all 23 quasars. *Upper panel:* Absolute luminosity. *Middle panel:* Luminosity of the 19 RQs normalized to unity at $\log \nu = 14.25$ (the longest wavelength available for all objects). Note that PG 1626+554 is the only quasar for which a steep α_x is associated with a strong soft excess; in all other objects a steep α_x is associated with a low 2 keV flux. Note also the two lowest spectra (PG 1001+054 and PG 1411+442), which appear to form a distinct group of “X-ray-weak” quasars. *Lower panel:* As in middle panel, for the four RLQs. PG 1425+267 may also be an “X-ray-weak” quasar.

also suggested by the presence of a marginally significant correlation between α_x and α_{ox} ($r_s = 0.533$, $\text{Pr} = 0.0089$; see below), and the absence of a significant correlation between α_x and α_{os} ($r_s = -0.201$, $\text{Pr} = 0.36$).

The X-ray luminosity distribution appears to be bimodal, with two quasars, PG 1001+054 and PG 1411+442, being a factor of 30 weaker than the mean radio-quiet quasar. These two quasars appear to form a distinct group of “X-ray-weak quasars.” The statistics for the radio-loud quasars (RLQs) are much poorer, and there is no well-defined mean, but PG 1425+267 may be a similar X-ray-weak RLQ.

3.3. Correlations with Emission-Line Properties

Figure 4 presents the correlations between the hard X-ray luminosity, $L_{2\text{ keV}}$ ($\log \nu = 17.685$) or the soft X-ray luminosity, $L_{0.3\text{ keV}}$, and the luminosity of $\text{H}\beta$, $[\text{O III}]$, He II , or Fe II . The value of r_s and the two-sided significance level (Pr) of r_s are indicated above each panel. Upper limits were not included in the correlations. Thus, the actual correlations for He II , where there are five upper limits, are likely to be smaller than found here (there is only one upper limit for $[\text{O III}]$ and Fe II , and none for $\text{H}\beta$). Excluding He II , the X-ray luminosity is most strongly correlated with $L_{\text{H}\beta}$ ($r_s = 0.734$, $\text{Pr} = 6.6 \times 10^{-5}$). We note in passing that $L_{\text{H}\beta}$ has an even stronger correlation with the luminosity at 3000 Å ($r_s = 0.866$, $\text{Pr} = 9 \times 10^{-8}$) and with the near-IR luminosity at 1 μm ($r_s = 0.810$, $\text{Pr} = 2.7 \times 10^{-6}$).

The position of the X-ray-weak quasars is marked in Figure 4. Both PG 1001+054 and PG 1411+442 appear to have an X-ray luminosity weaker by a factor of about 30 compared to other quasars with similar $L_{\text{H}\beta}$. PG 1425+267 is also weaker by a factor of ~ 10 compared with the other

RLQs. These ratios are the same as those found above in § 3.2, based on the spectral energy distribution.

Figures 5a–5d display various emission parameters that correlate with α_x (as obtained with $N_{\text{H}} = N_{\text{H1}}^{\text{Gal}}$). The FWHM of $\text{H}\beta$, $L_{[\text{O III}]}$, the $\text{Fe II}/\text{H}\beta$ flux ratio, and the ratio of $[\text{O III}]$ peak flux to $\text{H}\beta$ peak flux (as defined by Boroson & Green) are the emission-line parameters that correlate most strongly with α_x . As found in Paper I, all the α_x versus emission-line correlations become significantly weaker when we use α_x obtained with the free N_{H} fit.

The X-ray-weak quasars are labeled in the α_{ox} versus α_x correlation in Figure 5e. As expected, they have a steeper than expected α_{ox} for their α_x . The last parameter shown in Figure 5f is $1.5L_{14.25}^{1/2} \Delta v^{-2}$, where $\Delta v = \text{H}\beta$ FWHM. This parameter is related, under some assumptions, to the luminosity in Eddington units, as further discussed in § 4.7.

4. DISCUSSION

4.1. The Soft X-Ray Spectral Shape

We find an average spectral index $\langle \alpha_x \rangle = -1.62 \pm 0.09$ for the complete sample of 23 quasars, where the error here and below is the uncertainty in the mean. This slope is consistent with the mean slope $\langle \alpha_x \rangle = -1.57 \pm 0.06$ that we found for the subsample of 24 quasars out of the 58 AGNs analyzed by Walter & Fink (1993; the other 34 AGNs in their sample are Seyfert galaxies as defined by Véron-Cetty & Véron 1991). A similar average slope of -1.65 ± 0.07 was found by Schartel et al. (1996) for 72 quasars from the LBQS sample detected in the *ROSAT* all sky survey (RASS). Puchnarewicz et al. (1996) find a significantly flatter mean slope, $\langle \alpha_x \rangle = -1.07 \pm 0.06$, in a large sample of 108 soft X-ray-selected (0.5–2 keV) AGNs. Part of the difference is related to the exclusion of counts below 0.5 keV, which selects against steep α_x quasars, but this bias cannot explain the much flatter α_{ox} in their sample. As discussed by Puchnarewicz et al., their sample appears to include a large proportion of highly reddened quasars (see further discussion in § 4.2).

RLQs are known to have a flatter α_x than radio-quiet quasars (RQs) at energies above the PSPC band (e.g., Wilkes & Elvis 1987; Lawson et al. 1992). We find $\langle \alpha_x \rangle = -1.72 \pm 0.09$ for the 19 RQs, and $\langle \alpha_x \rangle = -1.15 \pm 0.14$ for the four RLQs in our sample. We find a similar trend using the Walter & Fink (1993) quasar data, where $\langle \alpha_x \rangle = -1.61 \pm 0.08$ for the RQs and $\langle \alpha_x \rangle = -1.36 \pm 0.08$ for the RLQs. A similar difference between RQs and RLQs was found by Ciliegi & Maccacaro (1996) in PSPC spectra of a sample of 63 AGNs extracted from the *Einstein* extended medium-sensitivity survey sample. We therefore conclude that the trend observed at harder X-rays also extends down to the 0.2–2 keV band. Schartel et al. (1996) stacked PSPC images of 147 RQs and 32 RLQs, finding for the sum images $\langle \alpha_x \rangle = -1.65 \pm 0.18$ for the RQs and $\langle \alpha_x \rangle = -1.00 \pm 0.28$ for the RLQs. However, the mean redshift of their objects is ~ 1.3 , and thus their results apply to the ~ 0.45 –4.5 keV band.

As discussed in Paper I, the *ROSAT* PSPC indicates a significantly different soft X-ray spectral shape for quasars compared with earlier results obtained by the *Einstein* imaging proportional counter (IPC) and *EXOSAT* low-energy and medium-energy (LE + ME) detectors (e.g., Wilkes & Elvis 1987; Masnou et al. 1992; Comastri et al. 1992; Saxton et al. 1993; Turner & Pounds 1989; Kruper,

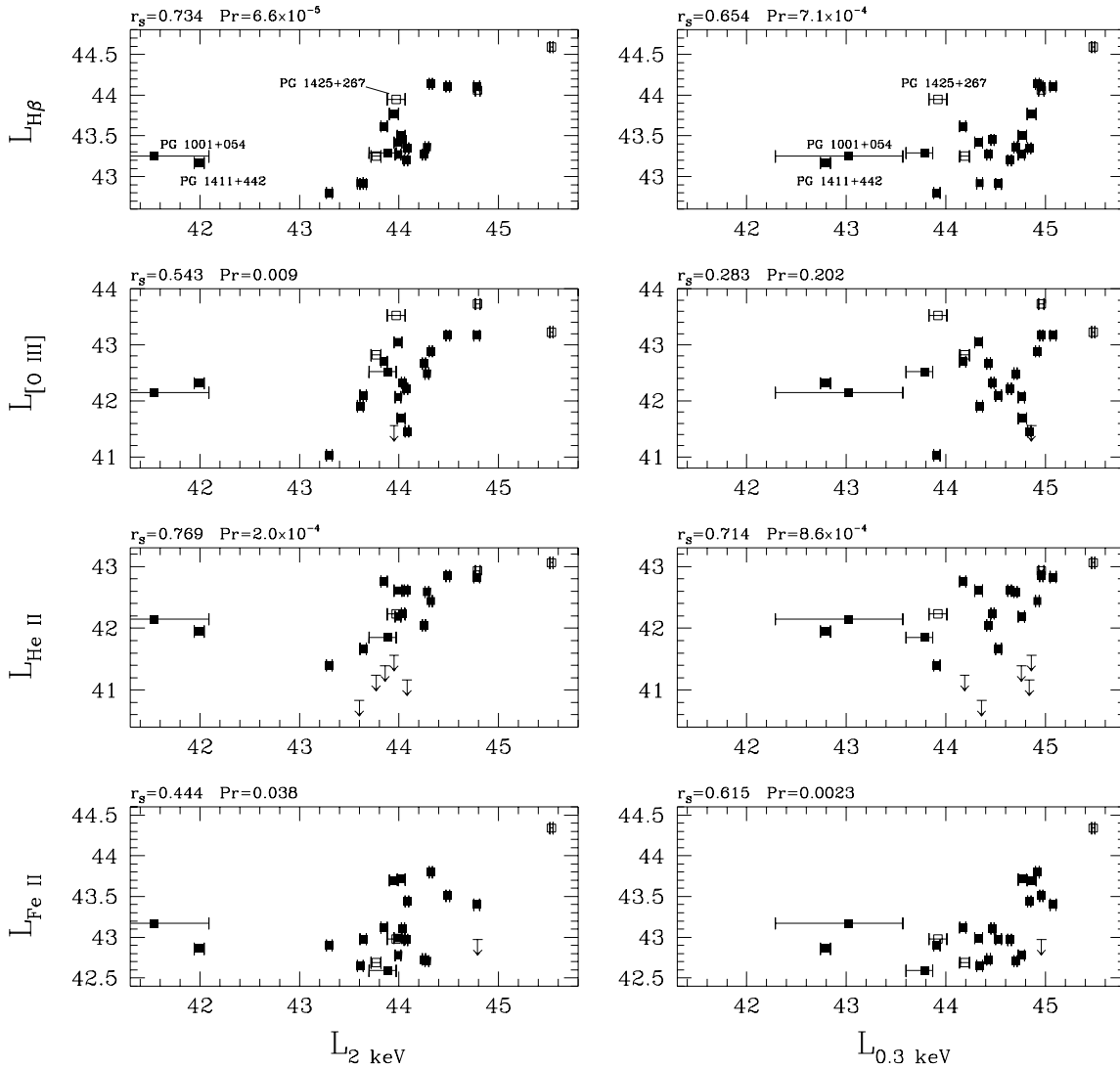


FIG. 4.—Correlations of line luminosity vs. X-ray luminosity. Open symbols are for radio-loud quasars. The Spearman rank-order correlation coefficient (r_s) and the two sided significance level (P), are indicated above each panel. The strongest correlation is displayed by the $H\beta$ luminosity (which correlates even better with the optical and near IR luminosity). Note the position of the X-ray-weak quasars. These appear to be fainter by factors of 10–30 in the X-ray than expected based on their $L_{H\beta}$. This factor is similar to the deficiency seen in Fig. 3.

Urry, & Canizares 1990). In particular, earlier missions suggested that the hard X-ray slope (Lawson et al. 1992; Williams et al. 1992) extends down to ~ 0.5 keV, with a steep rise at lower energy. Here we find that the 0.2–2 keV spectrum is fitted well by a single power law with Galactic absorption. This indicates that (1) the break between the soft and hard X-ray slope must occur well above 0.5 keV, (2) the break must be gradual, and (3) there is no steep soft component with significant flux down to ~ 0.2 keV. *ASCA* observations of two of the quasars in our sample, 3C 273 by Yaqoob et al. (1994) and PG 1116+215 by Nandra et al. (1996), find, as expected, significantly flatter spectra above 2 keV. However, the exact break energy cannot be accurately determined from the *ASCA* spectra because of likely calibration uncertainties below 1 keV.

As mentioned in Paper I, the different *Einstein* IPC and *EXOSAT* LE + ME results may be traced back to the combined effect of the lower sensitivity of these instruments below ~ 0.5 keV, and possibly some calibration errors. Small systematic errors in the PSPC response function appear to be present below 0.2 keV (Fiore et al. 1994), and

this instrument is thought to be significantly better calibrated at low energy than earlier instruments.

No significant spectral features are present in the PSPC spectra of any of 13 additional quasars reported here, indicating that intrinsic features must have an amplitude of less than 10%–20%. Note in particular the high S/N spectrum of PG 1116+215, where the number of counts is about 12 times the median sample counts, yet this spectrum is still consistent with a simple power law. For the complete sample we find that only one quasar, PG 1114+445, has a significant physical feature that is well described by a warm absorber model. In two other quasars, PG 1226+023 and PG 1512+370, there are significant features below 0.5 keV (Paper I). In the case of PG 1512+370 the features are at a level of $\sim 30\%$, and in PG 1226+023 they are at a level of $\lesssim 10\%$ and may well be due to small calibration errors. This result is consistent with the result of Fiore et al. (1994), who found that a simple power law provides an acceptable fit to the individual spectra of six high S/N PSPC quasar spectra.

A composite optical to hard X-ray spectral energy dis-

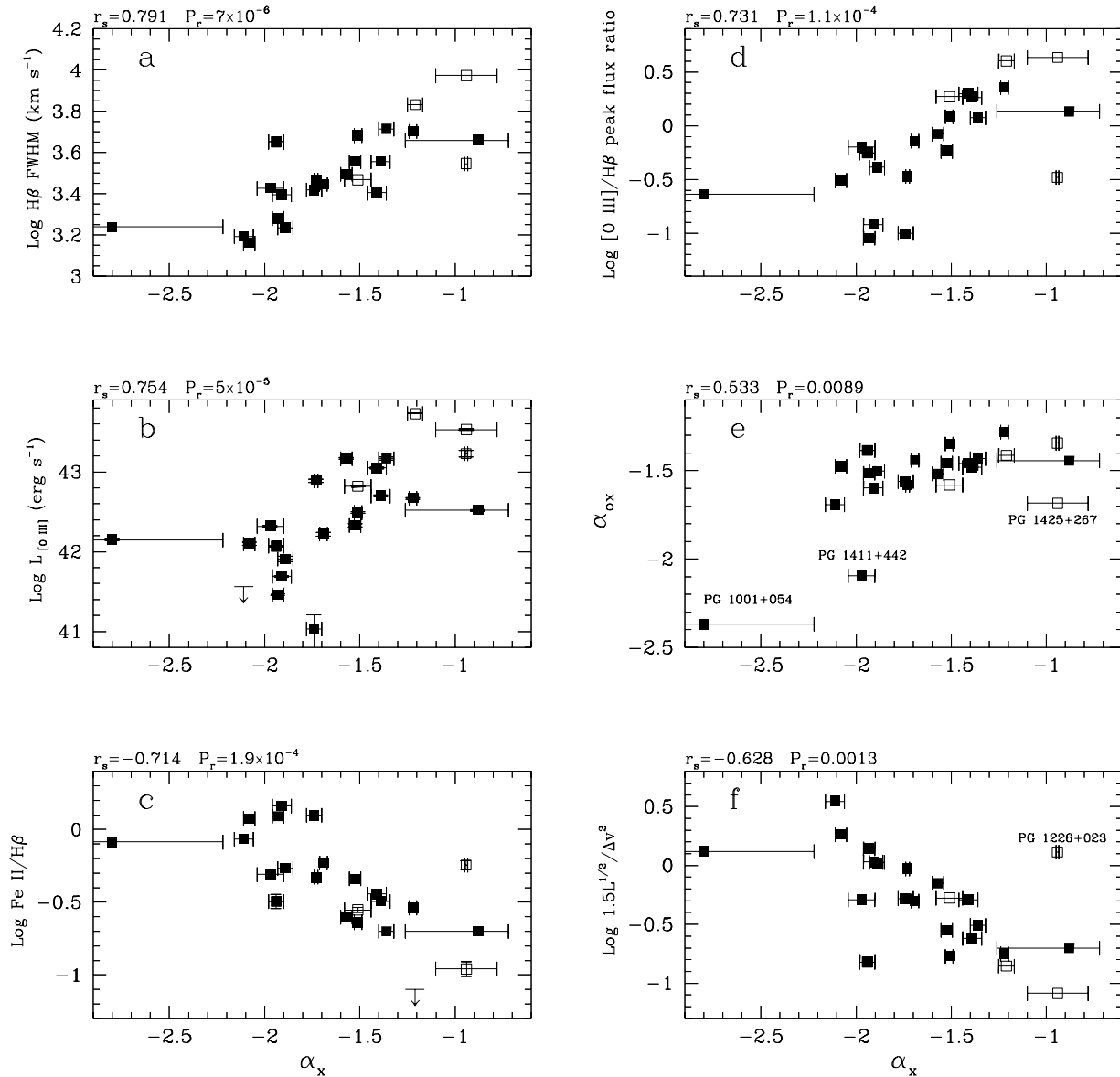


FIG. 5.—Various correlations with α_x . (a–d) Emission-line parameters that correlate strongly with α_x . Open symbols are for radio-loud quasars. The α_x versus H β FWHM correlation is the strongest in our sample. Note that the X-ray-weak quasars do not stand out in these correlations; they thus have the “right” spectral slope but the “wrong” luminosity. (e) Note the trend of increasing α_{ox} with increasing α_x , which indicates that a steep α_x tends to be associated with a weak 2 keV flux rather than a strong 0.3 keV flux. The position of the X-ray-weak quasars are marked. (f) The Y-axis is a measure of L/L_{Edd} under some simplified assumptions, as described in § 4.7.

tribution for RLQs and RQQs is displayed in Figure 6. To construct it, we used the mean $L_{14.25}$ (Table 5), the mean α_o , the mean α_x , and the mean α_{ox} in our sample. We excluded from the mean the three X-ray-weak quasars, and PG 1114+445, where α_x is highly uncertain because of the presence of a warm absorber. The mean spectra were extended above 2 keV assuming a slope of -1 for RQQs and -0.7 for RLQs. The Mathews & Ferland (1987, hereafter MF) quasar energy distribution is also displayed for the purpose of comparison. The MF shape assumes a steep soft component with a break to the hard X-ray slope above 0.3 keV, and it therefore significantly underestimates the soft X-ray flux at ~ 0.2 –1 keV.

RLQs tend to be somewhat stronger hard X-ray sources than RQQs. This trend, together with the flatter X-ray slope of RLQs, was interpreted by Wilkes & Elvis (1987) as possible evidence for a two-component model. In this interpre-

tation RLQs have the same hard X-ray component with $\alpha_x \sim -1$ as in RQQs, with an additional contribution from a flatter $\alpha_x \sim -0.5$ component, making their overall X-ray emission flatter and brighter. The additional X-ray component in RLQs could be related to the radio jet, e.g., through inverse Compton scattering. The composite spectrum suggests that although RLQs are brighter at 2 keV, they may actually be fainter at lower energy because of their flatter α_x . The RLQ composite is based only on four objects and is therefore rather uncertain. In addition, the results of Sanders et al. (1989, § IIIc) suggest that RLQs in the PG sample are about twice as bright at 2 keV compared with RQQs of similar optical luminosity, rather than the $\sim 30\%$ found for the composite; thus the difference in PSPC α_x would imply a smaller difference at 0.2 keV than is shown in the composite. If RLQs are indeed weaker than RQQs at 0.2 keV, then the two-component model suggested above

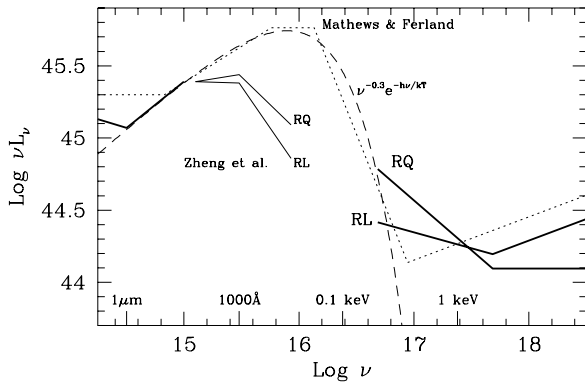


FIG. 6.—Composite optical-soft X-ray spectrum for the RQQs and RLQs in our sample (thick solid line). The three X-ray-weak quasars, and PG 1114+445, which is affected by a warm absorber, were excluded from the composite. The composite uses $\langle\alpha_{ox}\rangle = 1.482(\pm 0.025)$ and $\langle\alpha_x\rangle = 1.69(\pm 0.07)$ for the RQQs and $\langle\alpha_{ox}\rangle = 1.445(\pm 0.07)$ and $\langle\alpha_x\rangle = 1.22(\pm 0.16)$ for the RLQs. Note that despite the fact that RLQs are brighter at 2 keV, they are fainter at 0.3 keV. The Mathews & Ferland (1987) spectral shape assumes a hard X-ray power law down to 0.3 keV and a very steep component below 0.3 keV. This spectral shape is inconsistent with the PSPC results. The far-UV emission can be parameterized as a cutoff power law. The spectrum shown here has $T_{\text{cut}} = 5 \times 10^5$ K, which is close to the upper limit on T_{cut} . The Zheng et al. (1996) FUV composites for RLQs and RQQs are plotted in a thin solid line. They suggest that the FUV power law extends into the soft X-ray regime, with no extreme UV spectral break and no steep soft component below 0.2 keV.

would not be valid, and RLQs need to have a different X-ray emission process, rather than just an additional component.

The difference between RLQs and RQQs may actually be unrelated to the radio emission properties, as discussed in § 4.4.

Figure 6 also displays a simple cutoff power-law model of the form $L_\nu \propto \nu^{\alpha_o} e^{-h\nu/kT_{\text{cut}}}$ with $\alpha_o = -0.3$ and $T_{\text{cut}} = 5.4 \times 10^5$ K. This is an alternative way to interpolate between the UV and soft X-ray emission, and it is also a reasonable approximation for an optically thick thermal component. The lack of a very steep low-energy component down to 0.2 keV allows us to set an upper limit on T_{cut} . The upper limit is set using α_o and α_{os} , the slope from 3000 Å to rest-frame 0.15(1+z) keV (the lowest energy where the Galactic absorption correction error is $\leq 30\%$), given by

$$\alpha_{os} = [2.685\alpha_{ox} - (17.685 - \log \nu_s)\alpha_x] / \log(\nu_s/10^{15}),$$

where $\log \nu_s = 16.560 + \log(1+z)$. The upper limit on the cutoff temperature $T_{\text{cut}}^{\text{ul}}$ is related to the spectral slopes by

$$T_{\text{cut}}^{\text{ul}} = \frac{4.8 \times 10^{-11}(\nu_s - 10^{15}) \log e}{(\alpha_o - \alpha_{os}) \log(\nu_s/10^{15})} \text{ K}.$$

We find a rather small dispersion in $T_{\text{cut}}^{\text{ul}}$ with $\langle T_{\text{cut}}^{\text{ul}} \rangle = (5.5 \pm 2.6) \times 10^5$ K, averaged over the complete sample (Table 4), which corresponds to a cutoff energy of 47 eV, or about 3.5 rydbergs. This value of $T_{\text{cut}}^{\text{ul}}$ corresponds very closely to the far-UV continuum shape assumed by MF (see Fig. 6).

Walter et al. (1994) fitted such a cutoff model directly to six quasars and Seyfert galaxies, finding $\langle E_{\text{cut}} \rangle = 63 \pm 12$ eV, or $\langle T_{\text{cut}} \rangle = (7.3 \pm 1.4) \times 10^5$ K, while Rachen et al. (1996) find, using such a model, $\langle T_{\text{cut}} \rangle = (6.3 \pm 2.3) \times 10^5$ K for seven quasars and Seyfert galaxies. These values are consistent with our results. The small dispersion in T_{cut}

reflects the small dispersion in α_{os} in our sample, which is in marked contradiction to the dispersion predicted by thin accretion disk models, as further discussed in § 4.5.

4.1.1. The Far-UV Continuum

Zheng et al. (1996) have constructed a composite quasar spectrum based on *HST* spectra of 101 quasars at $z > 0.33$. They find a far-UV (FUV) slope (1050–350 Å) of $\langle\alpha_{\text{FUV}}\rangle = -1.77 \pm 0.03$ for RQQs and $\langle\alpha_{\text{FUV}}\rangle = -2.16 \pm 0.03$ for RLQs, with slopes of ~ -1 in the 2000–1050 Å regime. The Zheng et al. mean spectra, presented in Figure 6, together with the PSPC mean spectra, suggest that the FUV power-law continuum extends to the soft X-ray band. In the case of RQQs there is remarkable agreement in both slope and normalization of the soft X-ray and the FUV power-law continua. RLQs are predicted to be weaker than RQQs at ~ 100 eV by both the FUV and the PSPC composites. It thus appears that there is no extreme UV sharp cutoff in quasars, and that the fraction of bolometric luminosity in the FUV regime is significantly smaller than assumed.

The UV to X-ray continuum suggested in Figure 6 is very different from the one predicted by thin accretion disk models (§ 4.5) and suggested by photoionization models. In particular, it implies about a 4 times weaker FUV ionizing continuum compared with the MF continuum that was deduced based on the He II $\lambda 1640$ recombination line.

One should note, however, that the Zheng et al. sample is practically disjoint from our low- z sample, so it may still be possible that low- z quasars have a different FUV continuum.

4.2. Intrinsic Absorption

As shown in Figure 2, the H I column deduced from our accurate 21 cm measurements is consistent for all objects with the best-fit X-ray column. It is quite remarkable that even in our highest S/N spectra the 21 cm and X-ray columns agree to a level of about $1 \times 10^{19} \text{ cm}^{-2}$, or 5%–7%. This agreement is remarkable, since the 21 cm line and the PSPC are actually measuring the columns of different elements. Most of the soft X-ray absorption is due to He I or He II rather than H I, and the H I column is indirectly inferred assuming the column ratio H I/He I = 10. The fact that the 21 cm line and the PSPC give the same H I column implies that the H I/He column ratio at high Galactic latitudes must indeed be close to 10. The dispersion in the H I/He column ratio is lower than 20% (based on typical quasars in our sample), and may even be lower than 5% (based on our highest S/N spectra). There is therefore no appreciable Galactic column at high Galactic latitudes where the ionized fraction of H differs significantly from the ionized fraction of He, as found for example in H II regions (e.g., Osterbrock 1989).

The consistency of the 21 cm and X-ray columns also indicates that the typical column of cold gas intrinsic to the quasars in our sample must be smaller than the X-ray N_{H} uncertainty, or about $3 \times 10^{19}(1+z)^3 \text{ cm}^{-2}$. An additional indication for a lack of an intrinsic cold column in quasars comes from the fact that the strong correlations of α_x with the emission-line parameters described above (§ 3.3) become weaker when we use α_x from the free N_{H} fit rather than α_x from the fit with $N_{\text{H}}^{\text{Gal}}$. This indicates that $N_{\text{H}}^{\text{Gal}}$ is closer to the true N_{H} than the free N_{H} (see discussion in Paper I). In our highest S/N spectra we can set an upper limit of $\sim 2 \times 10^{19} \text{ cm}^{-2}$ on any intrinsic absorption. As discussed

in Paper I, the lack of intrinsic X-ray column for most quasars is consistent with more stringent upper limits set by the lack of a Lyman limit edge, as well as the He I and the He II bound-free edges in a few very high z quasars.

Puchnarewicz et al. (1996) suggest that the strong α_x versus α_{ox} correlation in their X-ray–selected sample is due to absorption of the optical and soft X-ray emission by cold gas and dust. They show that the α_x versus α_{ox} correlation for the 10 objects in Paper I can be explained by a universal spectral shape absorbed by a gas with a column of up to $N_H = 3 \times 10^{20} \text{ cm}^{-2}$ (see their Fig. 16). As described above, such absorbing columns are clearly ruled out by our high S/N spectra.

The mean α_o in the Puchnarewicz et al. (1996) sample is -0.92 ± 0.07 , which is significantly steeper than the mean α_o for optically selected quasars, e.g., a median of -0.2 for 105 PG quasars (Neugebauer et al. 1987), a median of -0.32 for 718 LBQS quasars (Francis et al. 1991), and $\langle \alpha_o \rangle = -0.36 \pm 0.05$ in our sample. Puchnarewicz et al. suggested that the much flatter α_o of the PG quasars is a selection bias, since these quasars were selected by the strength of their UV excess. However, the PG sample was selected on the basis of the color criterion $U - B < -0.44$, which, using the flux transformations of Allen (1973), corresponds to $\alpha_o \geq -1.8$. Thus, most of the red quasars discovered by Puchnarewicz et al. (1996) fit into the PG color criterion. The difference between the soft X-ray–selected and optically selected quasars must reflect the true tendency of quasars selected above 0.5 keV to be significantly redder than optically selected quasars. These red quasars may very well be affected by a large absorbing column ($N_H > 10^{21} \text{ cm}^{-2}$), as suggested by Puchnarewicz et al.

Intrinsic absorption is common in Seyfert 1 galaxies. About half of the primarily X-ray–selected Seyfert galaxies observed by Turner & Pounds (1989) using the *EXOSAT* LE + ME detectors, by Turner, George, & Mushotzky (1993a) using the *ROSAT* PSPC, and by Nandra & Pounds (1994) using the *Ginga* LAC for a largely overlapping sample, show low-energy absorption, or spectral features inconsistent with the simple power law typically observed above 2 keV. Quasars are very different. Excess absorption produces significant spectral features only in one object (PG 1114 + 445; see Paper I), i.e., $\sim 5\%$ (1/23) of the objects, and the absorbing gas is partially ionized (“warm”) rather than neutral. Given the typical S/N in our sample, we estimate that a partially ionized absorber that produces $\tau > 0.3$ can be ruled out in most of our objects. We cannot rule out partial absorption, or complete absorption and scattering, by a very high column density ($N_H > 10^{24} \text{ cm}^{-2}$) gas, since such effects may only suppress the flux level without affecting the spectral shape, and without inducing significant spectral features. As described in § 4.8, we suspect that such strong absorption may indeed be present in $\sim 10\%$ (3/23) of the quasars in our sample (the X-ray–weak quasars).

4.3. Implications of the Continuum-Continuum Correlations

The continuum-continuum luminosity correlations found here are all weaker than those found in Paper I. This is mostly due to the three X-ray–weak quasars that were not present in Paper I. For example, in Paper I we found that $f_{0.3 \text{ keV}}$ can be predicted to within a factor of 2, once $f_{1.69 \mu\text{m}}$ is given. This statement is still valid if the four extreme objects labeled in Figure 3 (*middle panel*) are excluded. The implications of the near-IR versus X-ray luminosity corre-

lation on the X-ray variability power spectrum were discussed in Paper I.

In Paper I we noted the similarity $\langle \alpha_{ox} \rangle = \langle \alpha_x \rangle = -1.50$, which was also noted by Brunner et al. (1992) and Turner et al. (1993a). However, we argued there that this similarity is only fortuitous, and that it does not imply that the X-ray power law can be extrapolated into the UV, since the optical slope is significantly different. Here we find that the relation $\langle \alpha_{ox} \rangle \simeq \langle \alpha_x \rangle$ holds only roughly for the complete sample where $\langle \alpha_{ox} \rangle = -1.55 \pm 0.24$ and $\langle \alpha_x \rangle = -1.62 \pm 0.45$. This relation does not hold when the sample is broken to the RQQs where $\langle \alpha_{ox} \rangle = -1.56 \pm 0.26$ and $\langle \alpha_x \rangle = -1.72 \pm 0.41$, and to the RLQs where $\langle \alpha_{ox} \rangle = -1.51 \pm 0.16$ and $\langle \alpha_x \rangle = -1.15 \pm 0.27$.

A significantly flatter $\langle \alpha_{ox} \rangle$ is obtained when the three X-ray–weak quasars and the absorbed quasar PG 1114 + 445 are excluded. Thus, “normal” RQQs in our sample have $\langle \alpha_{ox} \rangle = -1.48 \pm 0.10$, $\langle \alpha_x \rangle = -1.69 \pm 0.27$, while for the RLQs $\langle \alpha_{ox} \rangle = -1.44 \pm 0.12$, $\langle \alpha_x \rangle = -1.22 \pm 0.28$, where the \pm here and above refers to the dispersion about the mean rather than the error in the mean.

The α_x versus α_{ox} correlation found here is weaker than in Paper I, owing to the presence of the X-ray–weak quasars. However, the other 20 quasars appear to follow a trend of increasing α_x with increasing α_{ox} (Fig. 5e), indicating, as discussed in Paper I, that a steep α_x is generally associated with a weak hard X-ray component (at 2 keV) rather than a strong soft excess. The only object that clearly violates this trend is PG 1626 + 554 (Fig. 3), which has both a steep α_x and a strong soft excess. Puchnarewicz, Mason, & Córdova (1994) and Puchnarewicz et al. (1995a, 1995b) present PSPC spectra of three AGNs with extremely strong soft excess, where $L_{0.2 \text{ keV}} > L_{3000 \text{ \AA}}$. Our sample suggests that such objects are most likely rare, as can also be inferred from the selection criteria of Puchnarewicz et al. (1995a, 1995b), who selected their three objects from the *ROSAT* WFC all-sky survey, in which only five AGNs were detected. This selection criterion implies that these AGNs must have a very high far-UV flux.

The soft X-ray–selected quasars in the Puchnarewicz et al. (1996) sample have $\langle \alpha_{ox} \rangle = -1.14 \pm 0.02$, and none of their quasars is “X-ray weak,” i.e., with $\alpha_{ox} < -1.6$. The absence of X-ray–weak quasars in their sample is clearly a selection effect. The small survey area implies that most quasars in their sample are optically rather faint ($m_B \sim 18$ –19). “Normal” α_{ox} quasars in their sample produce a few hundred PSPC counts, but “X-ray–weak” quasars are below their detection threshold. The abundance of “X-ray–loud” quasars (i.e., $\alpha_{ox} > -1$) in the Puchnarewicz et al. sample is consistent with their rarity in optically selected samples. For example, quasars with $\alpha_{ox} = -1$ are about 20 times fainter at 3000 Å than quasars with $\alpha_{ox} = -1.5$, for the same L_x . Since the surface density of quasars increases as $\sim f_B^{-2.2}$ (see § 2.2.2.1 and Fig. 1 in Hartwick & Schade 1990), where f_B is the B -band flux, there are about 700 times more of these fainter quasars per B magnitude per square degree. Thus, even if only 0.3% of quasars at a given f_B have $\alpha_{ox} = -1$, there would still be twice as many quasars with $\alpha_{ox} = -1$ than $\alpha_{ox} = -1.5$ per square degree in an X-ray flux-limited sample, such as the Puchnarewicz et al. sample.

4.4. Implications of the Continuum-Line Correlations

The presence of the strong correlations of α_x with the H β FWHM, with $L_{[\text{O III}]}$, and with the Fe II/H β ratio described

in Paper I is verified here. The correlation coefficients for the complete sample are comparable to or somewhat smaller than those found in Paper I, but since the sample is larger, the significance level is now much higher (Fig. 5). We also report here an additional strong correlation of α_x with the ratio of [O III] peak flux to the H β peak flux, which is one of the emission-line parameters measured by Boroson & Green (1992). The α_x versus H β FWHM correlation is the strongest correlation we find between any of the X-ray continuum emission parameters and any of the emission-line parameters reported by Boroson & Green (with the addition of line luminosities reported in Table 5).

The α_x versus H β FWHM correlation is much stronger than the well-known α_x correlation with radio loudness ($r_s = 0.26$, Pr = 0.23 in our sample, but see Wilkes & Elvis 1987 and Shastri et al. 1993 for stronger correlations). Thus, the fact that the average α_x in RLQs is significantly flatter than in RQQs (§ 4.1) may be completely unrelated to the presence of radio emission, it may just reflect the fact that RLQs tend to have broader lines than RQQs (e.g., Tables 3 and 5 in Boroson & Green 1992). This appears to be the case in our sample, where the RLQs follow the same α_x versus H β FWHM distribution defined by the RQQs (Fig. 5a). This intriguing suggestion can be clearly tested by comparing α_x for RLQs and RQQs of similar H β FWHM.

Boller, Brandt, & Fink (1996) studied in detail narrow-line Seyfert 1 galaxies (NLS1), and they also find an apparently significant trend of increasing α_x with increasing H β FWHM. However, the scatter in their sample is significantly larger than that found here. In particular, they find a large range of α_x for H β FWHM < 2000 km s⁻¹, where only a few objects are available in our sample. The overall larger scatter in the Boller et al. (1996) data is probably due in part to the generally larger statistical errors in their α_x determinations. Large systematic errors may also be induced by the use of H β FWHM from a variety of sources. The measured H β FWHM can be sensitive to the exact measuring procedure, such as continuum placement, subtraction of Fe II blends, and subtraction of the narrow component of the line (produced in the narrow-line region), which may not be well resolved in low-resolution spectra. For example, Shastri et al. (1993) and Boroson & Green (1992) measured the H β FWHM independently for 13 overlapping objects, in eight of which their values differ by more than 1000 km s⁻¹. Aside from these technical reasons, the increased scatter may represent a real drop in the strength of the correlation when the luminosity decreases from the quasar level studied here to the Seyfert level studied by Boller et al. (1996). One should also note that intrinsic absorption is probably common in Seyfert 1 galaxies (Turner & Pounds 1989; Turner et al. 1993a), and such an absorption may lead to a large systematic error in α_x unless a high S/N spectrum is available indicating that features are present.

Wang et al. (1996) analyzed PSPC spectra of 86 AGNs, including 22 of the 23 quasars from our sample. Their sample is more heterogeneous than ours and includes some high- z quasars and a number of AGNs selected by their strong Fe II emission. The various correlations found by Wang et al. (1996) are typically similar to, or somewhat weaker than, those found here. For example, their (our) values are $r_s = -0.73$ (-0.79) for α_x versus H β FWHM and $r_s = 0.65$ (0.714) for α_x versus Fe II/H β . The somewhat smaller values found by Wang et al. may result from their

inclusion of $z > 0.4$ quasars, where α_x measures a higher energy slope than measured here, and from their use of a free N_H fit (limited from below by $N_{\text{HI}}^{\text{Gal}}$), which increases the random error in α_x (see Paper I).

We verify the strong correlation between $L_{\text{H}\beta}$ and $L_{2 \text{ keV}}$ found in Paper I. The correlations of the other lines with X-ray luminosity are significantly weaker than found in Paper I, and they are only marginally significant. Corbin (1993) found significant correlations of $L_{2 \text{ keV}}$ with Fe II/H β ($r_s = -0.474$), and of $L_{2 \text{ keV}}$ with the H β asymmetry ($r_s = -0.471$). We find that neither correlation is significant in our sample ($r_s = -0.288$, Pr = 0.19 and $r_s = -0.106$, Pr = 0.63). Since we can only test for correlation with $r_s > 0.64$, we cannot securely exclude the presence of the correlations reported by Corbin.

As discussed in Paper I, the α_x versus $L_{[\text{O III}]}$ correlation can be used to place a limit on the α_x variability on timescales shorter than a few years. Given the scatter in this correlation, we estimate that α_x should not vary by significantly more than 0.3 on these timescales.

It is hard to interpret the α_x versus [O III] to H β peak flux ratio correlation, since the physical meaning of the [O III] to H β peak flux ratio parameter defined by Boroson & Green (1992) is rather obscure. The [O III] peak flux is related to the width of [O III], and the [O III] to H β peak flux ratio may thus partly reflect the FWHM ratio of these lines. Thus, this correlation may represent a correlation of the [O III] FWHM with α_x . High spectral resolution measurements of the [O III] line profile are required to test this possibility.

4.5. Inconsistency with Thin Accretion Disk Models

Figure 7 presents the continuum emission from two thin accretion disk models. The models are for a disk around a rotating black hole, and viscous stress that scales as $(P_{\text{gas}} P_{\text{rad}})^{1/2}$, where P_{rad} is the radiation pressure and P_{gas} is the gas pressure (Laor & Netzer 1989). Significant soft X-ray emission is obtained for disks with a high accretion rate and a small inclination. However, as discussed by Fiore

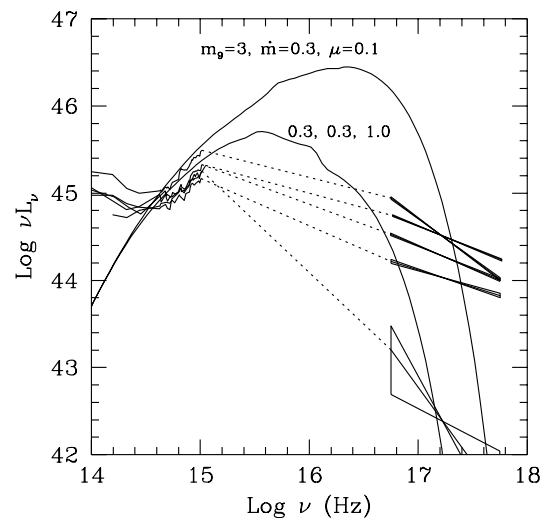


FIG. 7.—Observed energy distribution of some of the quasars in our sample vs. two accretion disk spectra. The disk models are for a rotating black hole, $\alpha = 0.1$, and viscous stress $\propto (P_{\text{gas}} P_{\text{rad}})^{1/2}$ (the other parameters are indicated). Note that although significant soft X-ray flux can be produced by the thin-disk models, the spectral slope is always much steeper than observed.

et al. (1995), the observed soft X-ray spectral slope is always much flatter than the one produced by a thin “bare” accretion disk model. As noted above, there is no indication in the 0.2–2 keV band for a very steep and soft “accretion disk” component.

Although thin disks cannot reproduce the 0.2–2 keV spectral shape, they may still be able to contribute a significant fraction of the flux at the lowest observed energy, i.e., 0.2–0.3 keV, above which a nonthermal power-law component sets in. As noted by Walter et al. (1994) and in Paper I, accretion disk models predict a large dispersion in the optical/soft X-ray flux ratio, and the strong correlation between these fluxes argues against the idea that a thin disk produces both the optical and soft X-ray emission. The arguments put by Walter et al. (1994) and in Paper I were only qualitative, and were not based on actual disk models. Furthermore, the objects in the small sample of Walter et al. were selected from known optically bright AGNs, and they also had to be bright soft X-ray sources, since most spectra were obtained from the *ROSAT* all-sky survey. Thus, these objects were a priori selected to be bright at both optical-UV and at soft X-rays, and the absence of a large scatter in the UV/soft X-ray flux ratio may just reflect the sample selection criteria. Such selection effects are not present in our sample, since the sample was defined independently of the X-ray properties, and X-ray spectra were obtained for all objects.

Below we describe a detailed calculation of the expected distribution of the optical/soft X-ray flux ratio for a complete optically selected sample based on the thin disk models of Laor & Netzer (1989), and show that such models cannot be reconciled with the observed distribution of the optical/soft X-ray flux ratio in our complete sample.

The optical/soft X-ray flux ratio, α_{os} , of a given disk model depends on the black hole mass, accretion rate \dot{m} , and inclination angle $\theta = \cos^{-1} \mu$. We now need to determine what distributions of these parameters will be consistent with the observed luminosity function in a complete optically selected sample. The intrinsic distribution of disk inclinations must be random. However, the observed distribution depends on the shape of the luminosity function of quasars, and possible obscuration effects, as described below.

The luminosity function of quasars is parameterized using the number density of quasars per unit volume per unit magnitude, $\Phi \equiv d^2N/dM dV$, and it is well fitted by a power law over a restricted range of magnitude M . Using Figure 2 in Hartwick & Schade (1990), we find $\log \Phi = 0.55M + c$ for $z < 0.2$ and $\log \Phi = 0.66M + c$ for $0.4 < z < 0.7$, where c is a constant. Since our sample is restricted to $z < 0.4$, we assume $\log \Phi = 0.6M + c$. Using the relation $M = -2.5 \log L + c$, we get $dn/dL \propto L^{-2.5}$, where $n \equiv dN/dV$.

The apparent luminosity of a flat disk, L_{app} , is related to its intrinsic luminosity through $L_{app} = 2\mu L$, neglecting limb-darkening effects that steepen the μ dependence and relativistic effects that flatten the μ dependence. This provides a reasonable approximation in the optical–near-UV regime (see Laor, Netzer, & Piran 1990). Assuming $dn/d\mu = \text{constant}$, i.e., a uniform distribution of inclination angles for the intrinsic quasar population, we would like to find $dn/d\mu$ for a given L_{app} . When L_{app} is fixed, $\mu \propto L^{-1}$, and substituting μ in the expression for dn/dL , we get $dn/d\mu \propto \mu^{0.5}$. Thus although the disks are assumed to have a

uniform distribution of inclination angles, the observed distribution at a given L_{app} is biased toward face-on disks.

To reproduce the observed luminosity function, we choose two values of \dot{m} , 0.1 and 0.3, where \dot{m} is measured in units of the Eddington accretion rate. Since $L \propto m_9 \dot{m}$, where m_9 is the black hole mass in units of $10^9 M_\odot$, the required mass distribution is $dn/dm_9 \propto m_9^{-2.5}$. The observed number of objects in a flux-limited sample is $dN_{ob}/dL \propto (dn/dL)V(L)$, where $V(L) \propto L^{3/2}$ is the observable volume for a flux-limited sample, such as the BQS sample. We therefore select a mass distribution of $dN_{ob}/dm_9 \propto m_9^{-1}$.

Figure 8 compares the observed distribution of α_{os} , as a function of νL_ν at 3000 Å, with the one expected from thin accretion disk models with the parameter distribution described above. Thin-disk models cannot account for the very small scatter in α_{os} .

The range of observed disk inclinations may actually be smaller than assumed here. For example, for a certain range of inclinations the disk may be completely obscured by an optically thick torus, as suggested in unification schemes for RQQs (e.g., Antonucci 1993). However, even if μ is fixed at a given value for all AGNs (say $\mu = 1$, which corresponds to the points extending from $\alpha_{os} = -1.5$ on the left axis to $\log \nu L_\nu = 46.5$ on the bottom axis of Fig. 8), the range in m_9 and \dot{m} will still produce a range in α_{os} which is much larger than observed.

The X-ray power-law emission is most likely produced by Comptonization of the thermal disk emission in a hot corona above the disk (e.g., Czerny & Elvis 1987). The slope and normalization of the power-law component are determined by the temperature and electron scattering optical depth in the corona (e.g., Sunyaev & Titarchuk 1985; Titarchuk & Lyubarskij 1995). The small range in α_{os} implies that some physical mechanism that couples the optical and soft X-ray emission processes must be operating, e.g., through a feedback that regulates both the temperature (see Haardt & Maraschi 1993) and the optical depth of the corona.

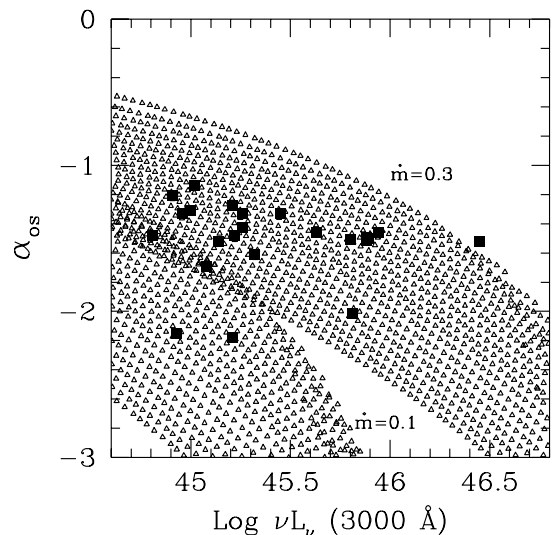


FIG. 8.—Observed α_{os} vs. luminosity distribution, and the one expected from a distribution of accretion disks that conforms to the observed luminosity function. The observed range in α_{os} is much smaller than the expected range, indicating that the UV and soft X-ray emission mechanisms are much better correlated than expected from thin-disk models.

As pointed out by various authors (Ross, Fabian, & Minshige 1993; Shimura & Takahara 1995; Dorner et al. 1996), and shown in Figure 7, simple thin accretion disks with no corona can produce a significant flux below 1 keV. For various disk model parameters α_{os} can in fact be significantly flatter than observed (Fig. 8), yet such extreme flat optical-soft X-ray spectra are only rarely observed (e.g., Puchnarewicz 1995b). The flattest spectra are expected for disks that are close to edge-on (e.g., Laor et al. 1990), and one therefore needs to assume that such disks are not observable. This is indeed expected in AGN unification schemes that invoke obscuring material close to the disk plane. Alternatively, the accreted material may form a geometrically thick, rather than a thin, configuration close to the center, which would display a smaller inclination dependence.

4.6. Inconsistency with Optically Thin Free-Free Emission Models

As was clearly demonstrated by Fiore et al. (1995) for six low-redshift quasars with high S/N PSPC spectra, and by Siemiginowska et al. (1995) using *Einstein* data for 47 quasars from Elvis et al. (1994), isothermal optically thin pure free-free emission models (Barvainis 1993) cannot fit the observed UV to soft X-ray energy distribution in AGNs. Furthermore, as was pointed out by Kriss (1995) and Hamann et al. (1995), optically thin free-free emission can also be ruled out based on the observed UV line emission.

4.7. On the Origin of the α_x versus H β FWHM Correlation

What is the physical process behind the α_x versus H β FWHM correlation? In Paper I we speculated that this may be either an inclination effect or, possibly, an L/L_{Edd} effect. Various authors raised the interesting suggestion that steep α_x quasars may be analogous to “high”-state Galactic black hole candidates (e.g., White, Fabian, & Mushotzky 1984; Fiore & Elvis 1995; Pounds, Done, & Osborne 1995), which display a steep slope in the soft and the hard X-ray bands when their brightness increases. The physical interpretation for this effect described by Pounds et al. (1995) is that the hard X-ray power law is produced by Comptonization in a hot corona and that, as the object becomes brighter in the optical-UV, Compton cooling of the corona increases, and the corona becomes colder, thus producing a steeper X-ray power law. This is obviously far from being a predictive model, since the coronal heating mechanism is not specified, and it is implicitly assumed that the coronal heating does not increase much as the quasar becomes brighter. However, the narrow H β line profiles provide independent evidence that steep α_x quasars may indeed have a higher L/L_{Edd} , as further described below.

The L/L_{Edd} of quasars can be estimated under two assumptions: (1) The bulk motion of the gas in the broad-line region is virialized, i.e., $\Delta v \simeq (GM/r)^{1/2}$, where $\Delta v = \text{H}\beta$ FWHM. This gives

$$\Delta v_{3000} = 2.19 m_9^{1/2} R_{0.1}^{-1/2},$$

where $\Delta v_{3000} = \Delta v/3000 \text{ km s}^{-1}$, $m_9 = M/10^9 M_\odot$, and $R_{0.1} = R/0.1 \text{ pc}$. (2) The size of the broad-line region is determined uniquely by the luminosity, $R_{0.1} = L_{46}^{1/2}$ pc, where $L_{46} = L_{\text{bolometric}}/10^{46}$. This scaling is consistent with reverberation line mapping of AGNs (Peterson 1993; Maoz 1995) and is theoretically expected if the gas in quasars is dusty (Laor & Draine 1993; Netzer & Laor 1994). Combin-

ing assumptions 1 and 2 gives

$$\Delta v_{3000} = 2.19 m_9^{1/2} L_{46}^{-1/4},$$

and thus the mass of the central black hole is

$$m_9 = 0.21 \Delta v_{3000}^2 L_{46}^{1/2}.$$

Using $L_{\text{Edd},46} = 12.5 m_9$, one gets

$$L/L_{\text{Edd}} = 0.38 \Delta v_{3000}^{-2} L_{46}^{1/2}.$$

Thus, given the two assumptions made above, narrow-line quasars should indeed have a high L/L_{Edd} , as previously suggested based only on their steep α_x and in analogy to Galactic black hole candidates.

To test whether L/L_{Edd} is indeed the underlying parameter that determines α_x , rather than just the H β FWHM, we looked at the correlation of α_x versus $\Delta v_{3000}^{-2} L_{46}^{1/2}$ displayed in Figure 5f, where we used the 1.7 μm luminosity and the relation $L_{\text{bolometric}} = 15 L_{14.25}$ (see Fig. 7 in Laor & Draine 1993). This correlation is not as strong as the α_x versus H β FWHM correlation, but it certainly appears suggestive. Note that $L/L_{\text{Edd}} > 1$ for some of the objects in the sample. These values are well above the thin accretion disk limit ($L/L_{\text{Edd}} = 0.3$; Laor & Netzer 1989) and suggest a thick-disk configuration. However, the assumptions used above to infer L/L_{Edd} are more qualitative than quantitative, since both the luminosity and the velocity field in the broad-line region may not be isotropic and therefore the presence of $L/L_{\text{Edd}} > 1$ cannot be securely deduced.

The Pounds et al. mechanism implies that a steep α_x is associated with a weak hard X-ray component, and, as described in § 3.2, this indeed appears to be the trend in our sample. If the Pounds et al. (1995) mechanism is true, then steep α_x AGNs should also have a steep hard X-ray power law. We are currently pursuing this line of research using *ASCA* and *SAX* observations of our sample.

An additional hint that a steep α_x may indeed be associated with a high L/L_{Edd} comes from the anecdotal evidence described by Brandt, Pounds, & Fink (1995), Brandt et al. (1994), Grupe et al. (1995), and Forster & Halpern (1996), where a number of Seyfert galaxies with a steep α_x display rapid, large-amplitude soft X-ray variability, which, as Boller et al. (1996) discuss, may imply a low-mass black hole and thus a high L/L_{Edd} . We are currently pursuing a more systematic study of the soft X-ray variability properties of broad-line versus narrow-line quasars using the *ROSAT* HRI.

The Pounds et al. (1995) suggestion is very appealing, since it allows a physical explanation for the tight correlation of apparently completely unrelated quantities. Although it is not clear a priori that α_x must steepen with increasing L/L_{Edd} , it appears that this is indeed what happens in Galactic black hole candidates.

Wang et al. (1996) also suggested that steep α_x objects have a high L/L_{Edd} , based on the fact that the fraction of luminosity emitted in the X-ray regime in thin accretion disk models increases with L/L_{Edd} , as discussed above in § 4.5. However, if this were indeed the physical process behind the α_x versus H β FWHM correlation, then one would expect high- L/L_{Edd} objects to have a high soft X-ray to optical flux ratio, while we find no correlation between H β FWHM and α_{os} ($r_s = -0.079$; Table 5).

We note in passing that one does not need to eliminate the normal broad-line region in narrow-line AGNs, as one of the options suggested by Boller et al. (1996) and Pounds

et al. (1995). The lines are narrow simply because of the lower black hole mass. The broad-line-emitting gas does not extend much closer to the center in narrow-line AGNs, as it does not extend much closer to the center in other AGNs, simply because of the effects of a higher ionization parameter, and a higher gas density, each of which quenches line emission.

4.8. The X-Ray-weak Quasars

Two of the quasars in our sample, the RQQs PG 1001+054 and PG 1411+442, and possibly also the RLQ PG 1425+267, appear to form a distinct group, which we term here “X-ray-weak” quasars, where the normalized X-ray luminosity is a factor of 10–30 smaller than the sample median. The position of these quasars as outliers can be noticed in the near-IR normalized flux distribution (Fig. 3), in the α_x versus α_{ox} correlation (Fig. 5e), and in the H β versus 2 keV and 0.3 keV luminosity correlations (Fig. 4). The first two indicators are based on the spectral shape, but the last one is independent of the spectral shape, and it also suggests a deficiency of the X-ray luminosity by a factor of 10–30 relative to the one expected based on the H β luminosity. An apparently bimodal distribution in α_{ox} can also be seen in Figure 5b of Wang et al. (1996), where six of their 86 quasars appear to form a distinct group with $\alpha_{ox} < -2$.

No bimodality of α_{ox} is seen in the Avni & Tananbaum (1986) *Einstein* study of the PG quasars. All the “X-ray-weak” quasars found by the *ROSAT* PSPC have $\alpha_{ox} < -2$, but none of the quasars detected by Tananbaum et al. (1986) have $\alpha_{ox} < -2$ (see Fig. 8 in Avni & Tananbaum 1986). The lack of $\alpha_{ox} < -2$ and bimodality in the Tananbaum et al. sample probably reflects its incompleteness, as only 86% of the quasars they observed were detected.

Although the three X-ray-weak quasars in our sample stand out in luminosity correlations, they conform well to the α_x correlations (Figs. 5a–5d). They thus have the “right” slope but the “wrong” flux level. Why are these quasars different? A simple answer is that for some unknown reason the X-ray emission mechanism, most likely Comptonization by $T \geq 10^8$ K electrons, tends to be bimodal, and in about 10% of quasars (or in all quasars for $\sim 10\%$ of the time) the X-ray flux level is strongly suppressed, while the spectral slope is not affected. Another option is that these are just normal quasars where the direct X-ray flux happens to be obscured. In this case what we see is only the scattered X-ray flux. Photoionization calculations indicate that a few percent of the direct flux will be scattered, depending on the covering factor of the absorber and the ionization parameter. If the ionization parameter is large enough, then the scattering will be mostly by free electrons, which preserves the spectral shape (see Netzer 1993 and Krolik & Kriss 1995). Such scattering will explain why the flux level is strongly reduced, while the spectral shape is not affected. Note that the obscuring matter should be transparent in the visible range, as is the case with the absorbing matter in broad absorption line QSOs.

Additional hints toward this interpretation come from the fact that PG 1411+442 is a broad absorption line quasar (BALQSO; Malkan, Green, & Hutchings 1987), and the UV absorbing gas may also produce soft X-ray absorption, as may also be the case in PHL 5200 (Mathur, Elvis, & Singh 1995). In addition, Green & Mathur (1996) find that BALQSOs observed by *ROSAT* have $\alpha_{ox} \lesssim -1.8$, i.e., as

observed here in the X-ray-weak RQQs. Another hint is provided by the fact that PG 1114+445 is also somewhat underluminous at 0.3 keV (Fig. 3b), and this quasar is most likely seen through a warm absorber. The X-ray-weak quasars could therefore be more extreme cases of PG 1114+445 and have an absorbing column that is large enough to completely absorb the direct soft X-ray emission.

Note that PG 1425+267 is a RLQ, while all BALQSOs are known to be RQQs (Stocke et al. 1992). It would thus seem implausible to suggest that PG 1425+267 is a BALQSO. However, PG 1425+267 has about the same relatively steep α_{ox} , compared to other RLQs, as in 3C 351 (Fiore et al. 1993), where X-ray absorption by warm gas is observed together with resonance UV absorption lines (Mathur et al. 1994) which are narrower than in “proper” BALQSOs.

Forthcoming *HST* spectra of all 23 quasars in our sample will allow us to test whether there is a one-to-one correspondence between X-ray weakness and broad absorption lines, i.e., whether all X-ray-weak quasars are BALQSOs, and not just that all BALQSOs are X-ray weak, as strongly suggested by the Green & Mathur (1996) and Green et al. (1996) results.

A simple test of whether these are truly “X-ray-weak quasars” or just normal highly absorbed quasars can be done by looking at their hard X-ray emission. If the X-ray column is below 10^{24} cm $^{-2}$, then the obscuring material would become transparent at $E < 10$ keV, and the observed hard X-ray emission will rise steeply above the cutoff energy, as seen in various highly absorbed AGNs, such as Mrk 3 (Iwasawa et al. 1994), NGC 5506 (Nandra & Pounds 1994), NGC 6552 (Fukazawa, et al. 1994; Reynolds et al. 1994), and NGC 7582 (Schachter et al. 1996). One may also expect significant spectral features produced by the obscuring material (e.g., Matt et al. 1996), depending on the ionization state of this material. Forthcoming *ASCA* observations of PG 1411+442 and PG 1425+267 will allow us to test this scenario.

Another prediction is that the X-ray-weak quasars should show lower variability compared with other quasars of similar X-ray luminosity. This is true for the following reasons: (1) They are intrinsically more X-ray luminous, and variability amplitude tends to drop with increasing luminosity (Barr & Mushotzky 1986; also Fig. 9 in Boller et al. 1996). (2) The scattering medium must be significantly larger than the X-ray source, and short-timescale variability will be averaged out. If the X-ray-weak quasars are just due to large-amplitude intrinsic variability of the soft X-ray emission, as seen in some steep-spectrum narrow-line Seyfert 1 galaxies (§ 4.7), then one may expect the exact opposite behavior, i.e., these quasars may become significantly brighter at soft X-rays at some stage in the future.

5. SUMMARY

We defined a complete sample of 23 optically selected quasars that includes all the PG quasars at $z \leq 0.400$, and $N_{\text{HI}}^{\text{Gal}} < 1.9 \times 10^{20}$ cm $^{-2}$. Pointed *ROSAT* PSPC observations were made for all quasars, yielding high S/N spectra for most objects. The high quality of the *ROSAT* spectra allows one to determine the best-fitting α_x with about an order of magnitude higher precision compared with previously available X-ray spectra. In this paper we report the observations of 13 quasars not described in Paper I, analyze the correlation of the X-ray properties of the complete

sample with other emission properties, determine the mean X-ray spectra of low- z quasars, and discuss the possible origin of the α_x versus $H\beta$ FWHM correlation, the nature of X-ray-weak quasars, and the physical origin of the soft X-ray emission. Our major results are the following:

1. The spectra of 22 of the 23 quasars are consistent, to within $\sim 10\%$ – 30% , with a single power-law model over the rest-frame range 0.2–2 keV. There is no evidence for significant soft excess emission with respect to the best-fit power law. We place a limit of $\sim 5 \times 10^{19} \text{ cm}^{-2}$ on the amount of excess foreground absorption by cold gas in most of our quasars. The limits are $\sim 1 \times 10^{19} \text{ cm}^{-2}$ in the two highest S/N spectra.

2. Significant X-ray absorption by partially ionized gas (“warm absorber”) in quasars is rather rare, occurring for $\lesssim 5\%$ of the population, which is in sharp contrast to lower luminosity AGNs, where significant absorption probably occurs for $\sim 50\%$ of the population.

3. The average soft X-ray spectral slope for RQQs is $\langle \alpha_x \rangle = -1.72 \pm 0.09$, and it agrees remarkably well with an extrapolation of the mean 1050–350 Å continuum recently determined by Zheng et al. (1996) for $z > 0.33$ quasars. For RLQs $\langle \alpha_x \rangle = -1.15 \pm 0.16$, which suggests that RLQs are weaker than RQQs below 0.2 keV, as suggested also by the Zheng et al. mean RLQ continuum. These results suggest that there is no steep soft component below 0.2 keV.

4. Extensive correlation analysis of the X-ray continuum emission parameters with optical emission-line parameters indicates that the strongest correlation is between α_x , and the $H\beta$ FWHM. A possible explanation for this remarkably strong correlation is a dependence of α_x on L/L_{Edd} , as observed in Galactic black hole candidates.

5. There appears to be a distinct class of “X-ray-weak” quasars, which form $\sim 10\%$ of the population, where the X-ray emission is smaller by a factor of 10–30 than expected based on their luminosity at other bands, and on their $H\beta$ luminosity.

6. Thin accretion disk models cannot reproduce the observed 0.2–2 keV spectral shape, and they also cannot reproduce the tight correlation between the optical and soft X-ray emission.

7. The H I/He I ratio in the ISM at high Galactic latitudes must be within 20%, and possibly within 5%, of the total H/He ratio.

The main questions raised by this study are the following:

1. What is the true nature of X-ray-quiet quasars? Are these quasars indeed intrinsically X-ray weak, or are they just highly absorbed but otherwise normal quasars?

2. What physical mechanism is maintaining the strong correlation between the optical–UV and the soft X-ray continuum emission, or, equivalently, maintaining a very small dispersion in the maximum possible far-UV cutoff temperature?

3. What is the physical origin for the strong correlations between α_x , and $L_{[\text{O III}]}$, Fe II/ $H\beta$, and the peak [O III]-to- $H\beta$ flux ratio?

4. Is the soft X-ray emission indeed related to the presence of radio emission, or is it just a spurious relation, and the primary effect is related to the $H\beta$ line width? Or, put differently, do RLQs and RQQs of similar $H\beta$ FWHM have similar α_x ?

Extensions of the *ROSAT* PSPC survey described in this paper to the hard X-ray regime with *ASCA* and *SAX*, and to the UV with *HST*, and soft X-ray variability monitoring with the *ROSAT* HRI, which are currently being carried out, may provide answers to some of the questions raised above. These studies will also allow us to (1) test whether steep α_x quasars have a steep 2–10 keV slope, as expected based on the Pounds et al. (1995) L/L_{Edd} interpretations; (2) test whether soft X-ray variability is indeed strongly tied to the $H\beta$ FWHM, as expected if the $H\beta$ FWHM is an indicator of L/L_{Edd} ; and (3) explore the relation of the UV line emission properties to the ionizing spectral shape.

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