Astronomy & Astrophysics manuscript no. arXiv_revised
July 8, 2016

The VIMOS Public Extragalactic Redshift Survey (VIPERS) ⋆ Measuring non-linear galaxy bias at $z \sim 0.8$

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Received –; accepted –

ABSTRACT

Aims. We use the first release of the VImos Public Extragalactic Redshift Survey of galaxies (VIPERS) of $\sim 50,000$ objects to measure the biasing relation between galaxies and mass in the redshift range $z = [0.5, 1.1]$.

Methods. We estimate the 1-point distribution function [PDF] of VIPERS galaxies from counts in cells and, assuming a model for the mass PDF, we infer their mean bias relation. The reconstruction of the bias relation is performed through a novel method that accounts for Poisson noise, redshift distortions, inhomogeneous sky coverage. and other selection effects. With this procedure we constrain galaxy bias and its deviations from linearity down to scales as small as 4 h^{-1} Mpc and out to $z = 1.1$.

Results. We detect small (up to 2 %) but statistically significant (up to 3 σ) deviations from linear bias. The mean biasing function is close to linear in regions above the mean density. The mean slope of the biasing relation is a proxy to the linear bias parameter. This slope increases with luminosity, which is in agreement with results of previous analyses. We detect a strong bias evolution only for $z > 0.9$, which is in agreement with some, but not all, previous studies. We also detect a significant increase of the bias with the scale, from 4 to 8 h^{-1} Mpc , now seen for the first time out to $z = 1$. The amplitude of non-linearity depends on redshift, luminosity, and scale, but no clear trend is detected. Owing to the large cosmic volume probed by VIPERS, we find that the mismatch between the previous estimates of bias at $z \sim 1$ from zCOSMOS and VVDS-Deep galaxy samples is fully accounted for by cosmic variance.

Conclusions. The results of our work confirm the importance of going beyond the over-simplistic linear bias hypothesis showing that non-linearities can be accurately measured through the applications of the appropriate statistical tools to existing datasets like VIPERS.

1. Introduction

Galaxies do not perfectly trace mass. The long known proof is that galaxy clustering depends on properties of galaxies such as luminosity, colour, morphology, stellar mass, and so on (e.g. Szapudi et al. 2000; Hawkins et al. 2001; Norberg et al. 2001, 2002; Zehavi et al. 2002, 2011; Meneux et al. 2009; Marulli et al. 2013) and not solely on the underlying mass distribution. Differences in clustering properties are caused by the physical processes that regulate the formation and evolution of galaxies and should disappear when averaging over scales much larger than those affected by these processes.

Modelling the physics of galaxy formation, or at least its impact on the *bias* relation, is of paramount importance to extract cosmological information from the spatial distribution of galaxies. Indeed, the large-scale structure of the Universe as traced by galaxies is one of the most powerful cosmological probes as testified by the increasing number of large galaxy redshift surveys either ongoing, such as Boss (Anderson et al. 2012), DES[1], and VIPERS (Guzzo et al. 2014) or those planned for the near future, such as eBOSS[2], DESI (Schlegel et al. 2011), and Euclid (Laureijs et al. 2011)[3]. These surveys are designed to address several important questions both in cosmology and in galaxy evolution theory. Chief among them is the origin of the accelerated expansion of the Universe.

It has recently been realised that geometry tests based on standard candles and standard rulers can trace the expansion history of the Universe but cannot identify the cause of the accelerated expansion, which can be obtained either by advocating a dark energy component or by modifying the gravity theory (e.g. Wang 2008). To break this degeneracy one needs independent observational tests. These are provided by the build-up of structures over cosmic time (Guzzo et al. 2008). The analysis of large-scale structures in galaxy distribution allows us to perform these two tests at one time. The baryonic acoustic oscillation peaks in the two point statistics provide a standard ruler to perform geometry test (e.g. Seo & Eisenstein 2003; Percival et al.

⋆ Based on observations collected at the European Southern Observatory, Paranal, Chile, under programmes 182.A-0886 (LP) at the Very Large Telescope, and also based on observations obtained with MegaPrime/MegaCam, a joint project of CFHT and CEA/DAPNIA, at the Canada-France-Hawaii Telescope (CFHT), which is operated by the National Research Council (NRC) of Canada, the Institut National des Science de l'Univers of the Centre National de la Recherche Scientifique (CNRS) of France, and the University of Hawaii. This work is based in part on data products produced at TERAPIX and the Canadian Astronomy Data Centre as part of the Canada-France-Hawaii Telescope Legacy Survey, a collaborative project of NRC and CNRS. The VIPERS web site is http://vipers.inaf.it/.

[1] www.darkenergysurvey.com
[2] http://www.sdss3.org/future/
[3] http://www.euclid-ec.org/

2007; Gaztañaga et al. 2009; Reid et al. 2012) whereas the apparent radial distortions in galaxy clustering caused by peculiar motions that are gravitationally induced allow us to measure the rate at which cosmic structures grow. Since both tests rely on baryonic structures, the knowledge of the bias relation is mandatory to probe the underlying mass distribution and set cosmological constraints. Notwithstanding, a clustering statistics that is in principle bias insensitive has been recently proposed by Bel & Marinoni (2014) and applied to VIPERS data (Bel et al. 2014).

Galaxy bias is not just a nuisance parameter in the quest for the world model. This bias also represents an opportunity to constrain models of galaxy evolution as it encodes important information about the physical processes that regulate the evolution of stars and galaxies. Therefore, it is important to model galaxy bias by establishing its link to the relevant astrophysical processes that regulate galaxy evolutions.In a recent review, Baugh (2013) has classified galaxy evolution models into two categories. The so-called empirical models belong to the first category. These authors use theoretically motivated relations to model galaxy distribution from halos extracted from N-body simulations. The two most popular schemes to populate halos with galaxies are halo occupation distribution (HOD; e.g. Cooray & Sheth 2002; Zheng et al. 2005) and sub-halo abundance matching (SHAM; e.g. Vale & Ostriker 2004; Conroy et al. 2006). The second category is represented by physical models in which the processes that regulate the evolution of baryons are explicitly considered to link them to the host dark matter structures. This approach is at the heart of the semi-analytic models of galaxy formation (e.g. White & Frenk 1991; Bower et al. 2006; De Lucia & Blaizot 2007). In most cases these models have been used to estimate galaxy bias from clustering statistics such as galaxy counts or 2-point correlation functions. The results indicate that the accuracy in both types of models is one of the main limitations in constraining dark energy or modified gravity from current and, even more so, future observational campaigns (Contreras et al. 2013).

Alternatively, one can adopt a purely phenomenological approach and use an operational definition of the bias in terms of map between the density fluctuations of mass, δ and galaxies, δ_g smoothed on the same scale. This approach assumes that galaxy bias is a local process that depends on the local mass density only. Many studies further assume that the bias relation is linear and deterministic, so that galaxy bias can be quantified by a single *linear bias* parameter b: $\delta_g = b\delta$. The concept of linear bias has played an important role in cosmology and many results have been obtained using this assumption, which is known to be unphysical as it allows negative densities. Also, this assumption has no justification at the relatively small scales of interest to the study of galaxy formation processes, which depend on many physical parameters and on large scales due to the presence of neutrinos (Villaescusa-Navarro et al. 2014). In fact, the bias is constant only on scales larger than about 40 h^{-1} Mpc (Manera & Gaztañaga 2011). Indeed, galaxy bias can be more conveniently described within a probabilistic framework as proposed by Dekel & Lahav (1999) and recently reformulated in the context of the halo model (Cacciato et al. 2012a).

From the phenomenological viewpoint, bias has been extensively investigated from counts in cells statistics, weak gravitational lensing, and galaxy clustering. The latter is probably most popular approach. It is typically based on 2-point statistics and on the assumption of linear bias (Norberg et al. 2001, 2002; Zehavi et al. 2005; Coil et al. 2006; Basilakos et al. 2007; Nuza et al. 2013; Arnalte-Mur et al. 2014; Skibba et al. 2014; Marulli et al. 2013). A comparatively smaller number of studies searched for deviations from the linear and deterministic bias either using 2-point (Tegmark & Bromley 1999) or higher order statistics (Verde et al. 2002; Gaztañaga et al. 2005; Kayo et al. 2004; Nishimichi et al. 2007; Swanson et al. 2008).

Gravitational lensing in the weak field regime has also been exploited to constrain galaxy bias. In particular, within the limit of scale-independent bias on large scales, weak lensing and galaxy clustering can be combined to estimate the linear bias parameter in a manner which is independent of the amplitude of density fluctuations (Amara et al. 2012; Pujol et al. 2016; Chang et al. 2016). On smaller scales weak lensing was also used to measure the scale dependence of galaxy bias (Hoekstra et al. 2002; Simon et al. 2007; Jullo et al. 2012; Comparat et al. 2013), although this effect is degenerate with bias stochasticity, i.e. the fact that galaxy bias might not be solely determined by the local mass density.

The most natural way to study a possible scale dependence (or non-linearity) of galaxy bias is in a probabilistic framework by means of counts in cells statistics (Sigad et al. 2000) since in this case one can separate deviations from linear bias and the presence of an intrinsic scatter in the bias relation. This approach was used to estimate the bias of galaxies in the PSCz (Branchini 2001), VVDS (Marinoni et al. 2005, hereafter M05), and zCOSMOS (Kovač et al. 2011, hereafter K11) catalogues as well as the relative bias of blue versus red galaxies in the 2 degrees field galaxy redshift survey (2dFGRS; Colless et al. 2001; Wild et al. 2005). Despite some disagreement, results obtained at low redshift ($z < 0.5$) generally indicate that, at least for some types of galaxies, the bias is stochastic, scale dependent and, therefore, non-linear. However. The situation at $z > 0.5$ is less clear. Gravitational lensing studies either focused on very bright objects to probe the baryonic acoustic oscillations (Comparat et al. 2013) or on galaxies in the COSMOS field (Jullo et al. 2012); these studies found no evidence for stochasticity but, in the case of Jullo et al. (2012), detected a significant scale dependence of galaxy bias. This conflicting evidence shows a lack of accuracy in current estimators for galaxy bias that is a serious warning for precision cosmology. This is especially true considering that this is the range that will be probed by next generation surveys that have the potential to trace both the redshift and scale dependence of galaxy bias (Di Porto et al. 2012a,b)

The results obtained so far that focus on counts in cells provide some conflicting evidence. In M05 authors analysed galaxies in the VVDS-Deep catalogue over an area of 0.4×0.4 deg and found significant deviations from linearity. The estimated effective linear bias parameter showed little evolution with redshift. In contrast, the biasing relation of zCOSMOS galaxies measured by K11 over a region of about 1.52 deg^2 turned out to be close to linear and rapidly evolving with the redshift. The tension between these results is paralleled by the observed differences in the spatial correlation properties of the two samples, with the 2-point correlation function in zCOSMOS systematically higher than that of VVDS galaxies (see e.g. Meneux et al. 2009). Owing to the large cosmic variance in the two samples, a rather small galaxy sample was proposed as the source of this mismatch, so a larger galaxy sample should be used to settle the issue.

The Vimos Public Extragalactic Redshift Survey [VIPERS] (Guzzo et al. 2014) has a depth similar to the zCOSMOS survey but with a much larger area of 24 deg^2. Its volume is comparable to that of 2dFGRS and is large enough to significantly reduce the impact of the cosmic variance (see Appendix in Fritz et al. 2014). We adopt the same approach as M05 and K11 and estimate galaxy bias from counts in cells. To do so we use a novel estimator that accounts for the effect of discrete sampling, allow-

ing us to use small cells and probe unprecedented small scales that are more affected by the physics of galaxy formation.

The layout of the paper is as follows. In Section 2 we describe both the real and mock datasets used in this work. In Section 3 we introduce the formalism used to characterise galaxy bias and the estimators used to measure this bias from a galaxy redshift survey. In Section 4 we assess the validity of the estimator and use mock galaxy catalogues to gauge random and systematic errors. We present our results in Section 5 and compare these with those of other analyses in Section 6. The main conclusions are drawn in Section 7

Throughout this paper we assume a flat ΛCDM universe (Ω_m, Ω_Λ, σ_8)= (0.25; 0.75; 0.9). Galaxy magnitudes are given in the AB system and, unless otherwise stated, computed assuming $h \equiv H_0/100\,\mathrm{km\,s^{-1}Mpc^{-1}} = 1$.The high value of σ_8 has little impact on our analysis since our results can be rescaled to different values of σ_8 that are more consistent with current cosmological constraints. The dependence of the magnitude upon h is expressed as $M = M_h - 5\log(h)$, where M_h is the absolute magnitude computed for a given h value.

2. Datasets

The results in this paper are based on the first release of the VIPERS galaxy catalogue (Garilli et al. 2014). Random and systematic errors were computed using a set of simulated galaxy catalogues mimicking the real catalogue and its observational selections. Both, the real and mock samples are described in this Section.

2.1. Real data

The VIMOS Public Extragalactic Redshift Survey is an ongoing ESA Large Programme aimed at measuring spectroscopic redshifts for about 10^5 galaxies at redshift $0.5 < z < 1.2$ and beyond. The galaxy target sample is selected from the 'T0005' release of the Canada-France-Hawaii Telescope Legacy Survey-Wide (CFHTLS-Wide) optical photometric catalogue[4]. VIPERS covers 24 deg^2 on the sky, divided over two areas within the W1 and W4 CFHTLS fields. Galaxies are selected to a limit of $I_{AB} < 22.5$, further applying a simple and robust colour preselection to efficiently remove galaxies at $z < 0.5$. This colour cut and the adopted observing strategy (Scodeggio et al. 2009) allow us to double the galaxy sampling rate with respect to a pure magnitude-limited sample. At the same time, the area and depth of the survey result in a relatively large volume, $5 \times 10^7\ h^{-3}$ Mpc^3, which is analogous to that of the 2dFGRS at $z \sim 0.1$. VIPERS spectra are collected with the VIMOS multi-object spectrograph (Le Fèvre et al. 2003) at moderate resolution (R = 210) using the LR Red grism, providing a wavelength coverage of 5500-9500 Å and a typical radial velocity error of $\sigma_v = 141(1+z)$ km s^{-1}.

The full VIPERS area of 24 deg^2 is covered through a mosaic of 288 VIMOS pointings. A complete description of the survey construction, from the definition of the target sample to the actual spectra and redshift measurements, is given in Guzzo et al. (2014). The dataset used in this and other papers of the early science release represent the VIPERS Public Data Release 1 (PDR-1) catalogue that includes 55359 redshifts (27935 in W1 and 27424 in W4), i.e. 64% of the final survey in terms of covered area (Garilli et al. 2014). A quality flag was assigned to each object in the process of determining their redshift from the spectrum, which quantifies the reliability of the measured redshifts. In this analysis, we use only galaxies with flags 2 to 9.5, which corresponds to a sample with a redshift confirmation rate of 90%.

Several observational effects need to be taken into account to investigate the spatial properties of the underlying population of galaxies.

i) Selection effects along the radial direction are driven by the flux limit nature of the survey and, at $z < 0.6$, by the colour preselection strategy. We use volume-limited (luminosity-complete) galaxy subsamples that we obtain by selecting galaxies brighter than a given magnitude threshold in a given redshift interval. We adopted a redshift-dependent luminosity cut of the form $M_B(z) = M_0 - z$ that should account for the luminosity evolution of galaxies (e.g. Zucca et al. 2009). The value of the threshold is set to guarantee that the selected sample is > 90 % complete within the given redshift interval. In this sense each subsample is volume limited and luminosity complete. This z-dependent luminosity cut is very popular and has been adopted in other papers (see e.g. K11). However, other works used different types of cuts, either ignoring any dependence on redshift (such as in M05; Coil et al. 2008) or assuming a different functional form for the redshift evolution (e.g. Arnalte-Mur et al. 2014). Adopting an incorrect luminosity evolution would generate a spurious radial gradient in the mean density of the objects and a wrong $z-$dependence in the galaxy bias. To minimise the impact of this potential bias, we carry out our analysis in relatively narrow redshift bins, so that adopting any of the aforementioned luminosity cuts would produce similar results, as we verified. The robustness of our result to the choice of the magnitude cut can be tested *a posteriori*. Figure 16 shows that the difference between estimates obtained with a z-dependent cut (filled red dot) and with a z-independent cut (open red dot) are smaller than the total random errors.

Selection effects induced by the colour preselection strategy were determined from the comparison between the spectroscopic and photometric samples (Guzzo et al. 2014; de la Torre et al. 2013; Fritz et al. 2014) and are accounted for by assigning to each galaxy an appropriate statistical weight dubbed colour sampling rate (CSR).

ii) The surveyed area presents regular gaps due to the specific footprint of the VIMOS spectrograph that creates a pattern of rectangular regions, called pointings, separated by gaps where no spectra are taken. Superimposed on this pattern are unobserved areas resulting from bright stars and technical and mechanical problems during observations. We discuss our strategy to take into account this effect in our counts in cells analysis in the following (see Cucciati et al. 2014, for a more detailed study).

iii) In each pointing, slits are assigned to a number of potential targets that meet the survey selection criteria (Bottini et al. 2005). Given the surface density of the targeted population, the multiplex capability of VIMOS, and the survey strategy, a fraction of about 45% of the parent photometric sample can be assigned to slits. We define the fraction of targets that have a measured spectrum as the target sampling rate (TSR) and the fraction of observed spectra with reliable redshift measurement as the spectroscopic sampling sate (SSR). Both functions are roughly independent of galaxy magnitude except the SSR, which decreases for $I_{AB} > 21.0$, as shown in Fig. 12 of Guzzo et al. (2014).

All these selection effects are thoroughly discussed and quantitatively assessed by de la Torre et al. (2013). We make no attempt to explicitly correct for these effects individually. In-

[4] http://terapix.iap.fr/cplt/oldSite/Descart/CFHTLS-T0005-Release.pdf

stead, we assess their impact on the estimate of galaxy bias in Section 4 using the mock galaxy catalogues described below.

For the scope of our analysis, the main advantages of VIPERS are the relatively dense sampling of tracers, which allows us to probe density fluctuations down to scales comparable to those affected by galaxy evolution processes, and the large volume that, as discussed in the previous Section, allows us to reduce the impact of cosmic variance considerably with respect to previous estimates of galaxy bias at $z \sim 1$.

The parent PDR-1 VIPERS sample contains 45871 galaxies with reliable redshift measurements. Here we restrict our analysis in the redshift range $z = [0.5, 1.1]$ since the number density of objects at larger distances is too small to permit a robust estimate of galaxy bias. To investigate the possible dependence of galaxy bias on luminosity and redshift, we partitioned the catalogue into subsamples by applying a series of cuts in both magnitude and redshift.

The complete list of subsamples considered in this work is presented in Table 1. We considered three redshift bins ($z = [0.5, 0.7]$, $[0.7, 0.9]$, $[0.9, 1.1]$) and applied different luminosity cuts that we obtained by compromising between the need of maximising both completeness and number of objects. Different luminosity cuts within each redshift bin allow us to study the luminosity dependence of galaxy bias at different redshifts. The magnitude cuts, $M_B = -19.5-z-5\log(h)$ and $-19.9-z-5\log(h)$, that run across the whole redshift range are used to investigate a possible evolution of galaxy bias. In the Table the subsamples are listed in groups. The first three groups indicate subsamples in the three redshift bins. The last group indicates subsamples that are designed to match the luminosity cuts performed by K11 ($M_B = -20.5 - z - 5\log(h = 0.7) = -19.72 - z - 5\log(h)$) and by M05 ($M_B = -20.0 - 5\log(h)$. The most conservative cut $M_B = -19.5 - z - 5\log(h)$ guarantees 90 % completeness out to $z = 1$ for the whole galaxy sample and higher for late type objects (see Fig. 1).

Since the analysis presented in this work is based on cell count statistics, a useful figure of merit is represented by the number of independent spheres that can be accommodated within the volume of the survey. Considering intermediate cells with a radius of 6 h^{-1} Mpc , the number of such independent cells is $N = 3869, 5527, 6964$ in the three redshift intervals $z = [0.5, 0.7]$, $[0.7, 0.9]$, $[0.9, 1.1]$, respectively.

2.2. Mock datasets

We considered a suite of mock galaxy catalogues mimicking the real PDR-1 VIPERS catalogue to assess our ability to measure the mean biasing function and evaluate random and systematic errors.

We used two different types of mock galaxy catalogues. We based the bulk of our error analysis on the first mock galaxy catalogue, which is described in detail in de la Torre et al. (2013). In this set of mocks, synthetic galaxies are obtained by applying the HOD technique to the dark matter halos extracted from the MultiDark N-body simulation (Prada et al. 2012)of a flat ΛCDM universe with (Ω_m, Ω_Λ, Ω_b, h, n, σ_8)= (0.27; 0.73; 0.0469; 0.7; 0.95; 0.82). Since the resolution of the parent simulation was too poor to simulate galaxies in the magnitude range sampled by VIPERS, de la Torre & Peacock (2013) applied an original technique to resample the halo field to generate sub-resolution halos down to a mass of $M = 10^{10}h^{-1}$ $M_\odot$. These halos were HOD populated with mock galaxies by tuning the free parameters to match the spatial 2-point correlation function of VIPERS galaxies (de la Torre et al. 2013). Once populated with HOD galaxies, the various outputs were rearranged to obtain 26 and 31 independent light cones mimicking the W1 and W4 fields of VIPERS and their geometry, respectively. In our analysis we considered 26 W1+W4 mock samples. They constitute our set of *Parent* mock catalogues, as opposed to the *Realistic* mock

Table 1. VIPERS subsamples.

z-range	M_B- cut $M_B - 5\log(h)$	n_{VIPERS} $10^{-3}h^3$ Mpc^{-3}	n_{mock} $10^{-3}h^3$ Mpc^{-3}
0.5 - 0.7	$-18.6 - z$	4.78	4.36
0.5 - 0.7	$-19.1 - z$	3.16	2.43
0.5 - 0.7	$-19.5 - z$	2.10	1.37
0.5 - 0.7	$-19.9 - z$	1.24	0.68
0.7 - 0.9	$-19.1 - z$	2.71	2.55
0.7 - 0.9	$-19.5 - z$	1.86	1.47
0.7 - 0.9	$-19.9 - z$	1.07	0.72
0.9 - 1.1	$-19.5 - z$	0.62	0.63
0.9 - 1.1	$-19.9 - z$	0.42	0.43
0.5 - 0.7	$-19.7 - z$	1.64	1.36
0.7 - 0.9	$-19.7 - z$	1.13	1.05
0.9 - 1.1	$-19.7 - z$	0.53	0.53
0.7 - 0.9	-20.0	1.42	1.49

Notes. Col. 1: redshift range. Col. 2: B-band magnitude cut (computed for $h = 1$). Col 3: galaxy number density in the real VIPERS sub-catalogues. Col 4: galaxy number density in the HOD-mock VIPERS sub-catalogues. In the *Parent* mock catalogue the number density is a factor $\sim$ 3.7 larger. Cells fully contained in the surveyed volume (i.e. not overlapping with gaps or empty areas) contain $\sim$ 40 % more objects on average.

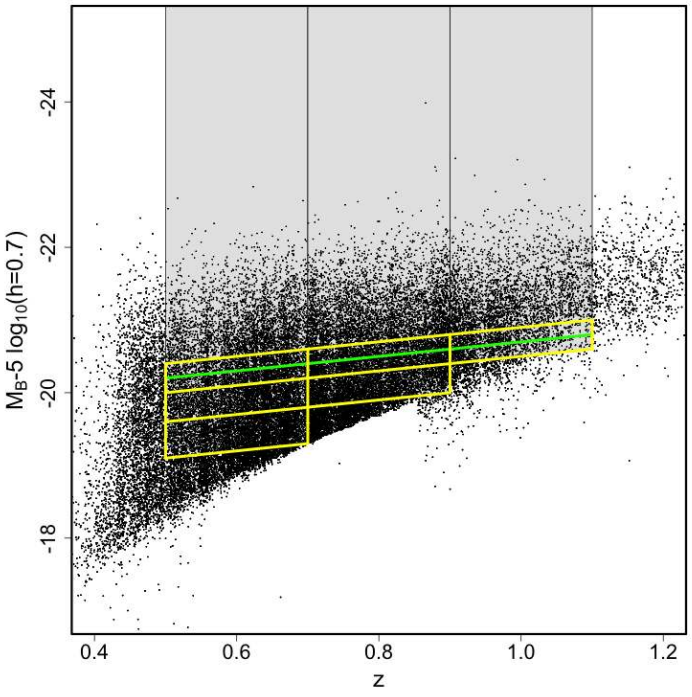

Fig. 1. Luminosity selection as a function of redshift. The black dots show the W1 and W4 VIPERS galaxies (with spectroscopic redshift flag between 2 and 9.5). Yellow lines represent the principal magnitude cuts applied in every redshift bin. The green line represents the cut $M_0 = -19.7 - z$ made to compare our results to those of K11.

catalogues that we obtain from the *Parent* set by applying the various selection effects (VIPERS footprint mask besides TSR, SSR, and CSR) and by adding Gaussian errors to the redshifts to mimic the random error in the measured spectroscopic redshifts. The mock catalogues were built assuming a constant SSR whereas, as we pointed out, this is a declining function of the apparent magnitude. However, the dependence is weak and only affects faint objects, i.e. preferentially objects at large redshifts. For this reason we decided to explicitly include this dependence by selectively removing objects, starting from the faintest and moving towards brighter objects until we match the observed SSR(m) (Guzzo et al. 2014).

The average galaxy number densities in the mocks are listed in Column 4 of Table 1. For $z \leq 0.9$ the number density in the mocks is similar or somewhat smaller than in the real catalogue. This discrepancy increases with the luminosity and probably originates from the uncertainty in the procedure to HOD-populate halos with bright mock galaxies. The consequence for our analysis is an overestimation of the random errors in the measurement of the bias of VIPERS galaxies. At higher redshift the trend is reversed; the number density of objects in the mocks is systematically larger than in the real catalogue. In this case, to avoid underestimating errors, we randomly diluted the galaxies in the mocks. Hence the perfect match of number densities in the redshift bin $z = [0.9, 1.1]$, as shown in the table.

On the smallest scale investigated in this paper, $R = 4\ h^{-1}\,\mathrm{Mpc}$, the second-order statistics of simulated galaxies and the variance of the galaxy density field are underestimated by ~ 10 % (Bel et al. 2014). Therefore, to check the robustness of our bias estimate to the galaxy model used to generate the mock catalogues and to the underlying cosmological model, we considered a second set of mocks. These were obtained from the Millennium N-body simulation (Springel et al. 2005) of a flat ΛCDM universe with $(\Omega_m, \Omega_\Lambda, \Omega_b, h, n, \sigma_8) = (0.25; 0.75; 0.045; 0.73; 1.00; 0.9)$ and using the semi-analytic technique of De Lucia & Blaizot (2007), an alternative to the HOD. As a result of the limited size of the computational box, it was possible to create light cones with an angular size of 7×1 deg^2, i.e. smaller than the individual W1 and W4 fields. Overall, we considered $26 + 26$ reduced versions of the W1+W4 fields. From these light cones we created a corresponding number of *Realistic* mock catalogues.

Robustness tests that involve both types of mock catalogues were restricted to a limited number of samples (one for each redshift bin). In these tests we simply compared the errors in the bias estimates after accounting for the larger cosmic variance in the Millennium mocks. Since these robustness tests turned out to be successful in the sense that errors estimated with the two sets of mocks turned out to be consistent with each other, we do not mention these mocks again and, for the rest of the paper, fully rely on the error estimates obtained from the HOD mocks.

3. Theoretical background

In this section we briefly describe the formalism proposed by Dekel & Lahav (1999) and the method that we use to estimate bias from galaxy counts. The key step is the procedure to estimate the galaxy PDF, $P(\delta_g)$, from the measured probability of galaxy counts in cells, $P(N_g)$. We review some of the techniques proposed to perform this crucial step and describe in detail the technique used in this work.

3.1. Stochastic non-linear bias

Dekel & Lahav (1999) proposed a probabilistic approach to galaxy bias in which non-linearity and stochasticity are treated independently. In this framework, galaxy bias is described by the conditional probability of galaxy over-density, δ_g, given the mass over-density δ: $P(\delta_g|\delta)$. Both quantities are smoothed on the same scale and treated as random fields. If biasing is a local process then $P(\delta_g|\delta)$ fully characterises galaxy bias. Key quantities formed from the conditional probability are the mean biasing function

$$b(\delta)\delta \equiv \langle \delta_g|\delta \rangle = \int P(\delta_g|\delta)\delta_g d\delta_g \ , \tag{1}$$

and its non-trivial second-order moments

$$\hat{b} \equiv \frac{\langle b(\delta)\delta^2 \rangle}{\sigma^2} \qquad \tilde{b}^2 \equiv \frac{\langle (b(\delta)\delta)^2 \rangle}{\sigma^2} \ , \tag{2}$$

where $\sigma^2 \equiv \langle \delta^2 \rangle$ is the variance of the mass over-density field on the scale of smoothing. The quantity $\hat{b}$ represents the slope of the linear regression of δ_g against δ and is the natural generalisation of the linear bias parameter. The ratio $\tilde{b}/\hat{b}$ quantifies the deviation of the mean biasing function from a straight line. It measures the non-linearity of the mean biasing relation and, in realistic cases, is close to unity. In the limit of linear and deterministic bias, the two moments $\hat{b}$ and $\tilde{b}$ coincide with the (constant) mean biasing function $b(\delta) = b_{\rm LIN}$, where $b_{\rm LIN}$ is the familiar linear bias parameter. We note that $\hat{b}$ is sensitive to the mass variance and scales as $\hat{b} \propto \sigma^{-1}$. On the contrary, the moments' ratio is very insensitive to it, $\tilde{b}/\hat{b} \propto \sigma^{0.15}$ (Sigad et al. 2000). These scaling relations are used in Section 5 to compare results obtained assuming different values of σ_8. There are other useful parameters related to galaxy bias that can be measured from the data. One is the ratio of variances $b_{\rm var} \equiv (\sigma_g/\sigma)^2$ in which σ_g is measured from counts in cells and σ depends on the assumed cosmological model. Another quantity is the inverse regression of δ over δ_g, $b_{\rm inv} \equiv \sigma_g^2/\langle \delta_g \delta \rangle$ that requires an estimate of the galaxy and the mass density fields (Sigad et al. 1998). In the case of non-linear deterministic bias these quantities differ from $\hat{b}$. Specifically, if the non-linearity parameter $\tilde{b}/\hat{b}$ is larger (smaller) than unity then they are biased high (low) with respect to $\hat{b}$ (Dekel & Lahav 1999).

In this paper we focus on the $\hat{b}$ parameter, a choice that allows us to compare our results with those of K11 (but not with M05, in which the focus is instead on $\tilde{b}$). Fortunately, as we shall see, the small degree of non-linearity makes these two choices almost equivalent.

If bias is deterministic, then it is fully characterised by the mean biasing function $b(\delta)\delta$. However, we do not expect this to be the case since galaxy formation and evolution are regulated by complex physical processes that are not solely determined by the local mass density. Therefore, for a given value of δ there is a whole distribution of δ_g about the mean $b(\delta)\delta$. This scatter, often referred to as bias stochasticity, is contributed by two sources: shot noise due to the discrete sampling of a continuous underlying density field and those astrophysical processes relevant to the formation and evolution of galaxies that do not depend (solely) on the local mass density.

Previous studies (Branchini 2001; Marinoni et al. 2005; Viel et al. 2005; Kovač et al. 2011) that, like this one, used the galaxy 1-point PDF to recover the biasing function ignored the impact of stochasticity and assumed a deterministic bias. We aim to improve the accuracy of the bias estimator by taking bias stochasticity into account and we do this by assuming that shot noise is

the only source of stochasticity. This simplifying assumption can be justified theoretically by both numerical and analytic arguments. Numerical experiments in which semi-analytic galaxies are used to probe the mass density field in samples mimicking SDSS (Szapudi & Pan 2004) (see Figs. 11 and 16) and 2MRS (Nusser et al. 2014) (see Fig. 1), i.e. two surveys with galaxy number densities similar to that of VIPERS, do indeed show that shot noise is the dominant source of scatter. More specifically, Poisson noise accounts for the scatter in the δ_g versus δ relation except at large over-density where the relation is over-dispersed. Analytic arguments in the framework of the halo model also confirm that the main source of stochasticity is shot noise with the halo-halo scatter providing a significant contribution for faint objects alone (Cacciato et al. 2012b). Assessing the impact of this shot noise only assumption is not simple, but some arguments can be made to quantify the systematic effect of underestimating stochasticity.

An upper limit can be obtained when stochasticity is ignored altogether. In the case of linear and stochastic bias, for example, $\tilde{b}$ and $\hat{b}$ would be equal whereas $b_{\rm inv}$ would be systematically larger by about 10% (Somerville et al. 2000). The more realistic case of a non-linear and stochastic bias was considered by Sigad et al. (2000) using numerical simulations again. In this case, the effect of ignoring stochasticity is that of overestimating both $\tilde{b}$ and $\hat{b}$. The amplitude of the effect depends on both the cosmological model assumed and the scale considered. To obtain estimates relevant to our analysis we repeated the Sigad et al. (2000) test in Section 4.1. The results, which we anticipate here, indicate that $\tilde{b}$ and $\hat{b}$ are overestimated by 8(4)% on a scale of 4(8) h^{-1} Mpc . As for the ratio, $\tilde{b}/\hat{b}$ we also confirm that it is remarkably insensitive to stochasticity and, as expected, to the model adopted (Sigad et al. 2000).

Analyses of the datasets may also constrain the size of the effect. Galaxy clustering, higher order statistics, or gravitational lensing generally indicate that galaxy bias cannot be linear and deterministic. However, as we anticipated in the introduction, it is not possible to disentangle the effects induced by non-linearity and stochasticity, except for the case of relative bias between two types of tracers. With respect to this, the largest stochasticity $\sigma_b/\hat{b} = 0.44$ so far was measured by Wild et al. (2005). If ignored, this would induce a systematic error of $\sim 20\%$ on the relative $\hat{b}$ moments.

Overall, the variety of evidenceindicates that if stochasticity is ignored then σ_b and $\hat{b}$ are overestimated by 10-20 %, whereas their ratio is unaffected. However, we stress that in our work stochasticity is, at least in large part, taken into account. Therefore, we expect that our assumption that shot noise is the only source of bias stochasticity generates systematic errors well below the 10 % level.

3.2. Direct estimate of $b(\delta)\delta$

Under the hypothesis that bias is deterministic and monotonic the mean biasing function, $b(\delta)\delta$, can be estimated by comparing the PDFs of the mass and of the galaxy over-density. We let $C(\delta) \equiv P(>\delta)$ and $C_g(\delta_g) \equiv P(>\delta_g)$ be the cumulative probability distribution functions [CDFs] obtained by integrating the two PDFs. Monotonicity guarantees that the ranking of the fluctuations δ and δ_g is preserved and $b(\delta)\delta$ can be obtained by equating the two CDFs at the same percentile,

$$b(\delta)\delta = C_g^{-1}(C(\delta)) \,, \tag{3}$$

where C_g^{-1} indicates the inverse function of C_g.

Equation 3 provides a practical recipe to estimate galaxy bias from observed counts in cells of a given size. It requires three ingredients: the galaxy over-density δ_g, its PDF, and that of δ. δ_g can be estimated from galaxy counts in cell, N_g as

$$1 + \delta_g = N_g/\langle N_g \rangle \,, \tag{4}$$

where $\langle N_g \rangle$ represents mean over all counts. From Eq. 4 one can form the galaxy PDF, $P(\delta_g)$ and the count probability $P(N_g)$. The biasing function can then be obtained by comparing $C_g(\delta_g)$ with a model $C(\delta)$.

This simple bias estimator has been used by several authors (Sigad et al. 2000; Branchini 2001; Marinoni et al. 2005; Viel et al. 2005; Kovač et al. 2011). It is potentially affected by several error sources that should be systematically investigated. The first error source is shot noise that affects the estimate of δ_g from N_g. Shot noise induces stochasticity in the bias relation in contrast with the hypothesis of deterministic bias. Stochasticity affects the estimate of $b(\delta)\delta$ from Equation 3, especially at large values of δ_g, where the CDF flattens and the evaluation of the inverse function C_g^{-1} becomes noisy. A second issue is the mass PDF for which no simple theoretical model is available. The last error source is redshift distortions. Galaxy over-densities are computed using the redshift of the objects rather than distances. This induces systematic differences between densities evaluated in real and redshift space (Kaiser 1987).

All these issues potentially affect the estimate of galaxy bias and should be properly quantified and accounted for. In the next section, we review some existing estimators designed to minimise the impact of the shot noise and propose a new estimator that we apply in this paper. We investigate the performance of this new strategy in Section 4.

3.3. From $P(N_g)$ to $P(\delta_g)$...

The probability of galaxy counts, $P(N_g)$, can be expressed as

$$P(N_g) = \int_{-1}^{+\infty} P(\delta_g) P(N_g|\delta_g) d\delta_g \,, \tag{5}$$

where the conditional probability function $P(N_g|\delta_g)$ specifies the way in which discrete galaxies sample the underlying, continuous field. The common assumption that galaxies are a local Poisson process implies that

$$P(N_g|\delta_g) = \frac{\left[\langle N_g \rangle (1+\delta_g)\right]^{N_g} e^{-\langle N_g \rangle (1+\delta_g)}}{N_g!} \,. \tag{6}$$

The Poisson model provides a good match to numerical experiments except at large densities where a negative binomial distribution seems to provide a better fit (Sheth 1995; Somerville et al. 2001; Casas-Miranda et al. 2002). In this work we adopt the Poisson model. However, different forms for $P(N_g|\delta_g)$ could be considered as well.

The following strategies have been proposed to estimate $P(\delta_g)$ from $P(N_g)$ using Equation 5:

- *Richardson-Lucy deconvolution*. Szapudi & Pan (2004) proposed this iterative, non-parametric method to reconstruct $P(\delta_g)$ by comparing the observed $P(N_g)$ to that computed from Eq. 5 at each step of the iteration, starting from an initial guess for $P(\delta_g)$.
- *Skewed lognormal model fit*. This parametric method was also implemented by Szapudi & Pan (2004). In this approach one assumes a skewed lognormal form for $P(\delta_g)$ and then determines the four free parameters of the model by minimising the difference between Eq. 5 and the observed $P(N_g)$.

– *Gamma expansion* [Γ_E]. Among the various forms proposed to model the galaxy PDF, the Gamma expansion, defined by expanding the Gamma distribution on a basis of Laguerre polynomials (Mustapha & Dimitrakopoulos 2010) captures the essential features of the galaxy density field. The expansion coefficients directly depend on the moments of the observed counts. Because of this, the full shape of the galaxy PDF can be recovered directly from the observed $P(N_g)$ with no need to integrate Eq. 5.

Szapudi & Pan (2004) have tested the ability of the first two methods in reconstructing the PDF of halos and mock galaxies obtained from N-body simulations. They showed that a successful reconstruction can be obtained when the sampling is $\langle N_g \rangle \geq 0.1$; safely a factor 3 smaller than the smallest mean galaxy density in our VIPERS subsamples. Bel et al. (2016) extensively tested the Γ_E-method and showed, using the same mock catalogues as in this paper, that this method reconstructs the PDF of a VIPERS-like galaxy distribution with an accuracy that is superior to that of the other methods. This comes at the price of discarding counts in cells that overlap the observed areas by less than 60 %, which is a constraint that further reduces deviations from the Poisson sampling hypothesis.

To illustrate the performance of the Γ_E-method we plot, in Figure 2, the galaxy PDFs Γ_E-reconstructed from the 26 *Realistic* mock VIPERS subsamples with galaxies brighter than $M_B = -19.1 - z - 5log(h)$ in the range $z = [0.7, 0.9]$. The blue dashed curve represents the mean among the mocks and the blue band the 1-σ scatter. The scatter for cells of $R = 8\ h^{-1}$ Mpc is larger than for $R = 4\ h^{-1}$ Mpc and is driven by the limited number of independent cells rather than sparse sampling.

The reconstruction is compared with the "reference" PDF (solid, red line) obtained by averaging over the PDFs reconstructed, with the same Γ_E method, from the *Parent* mock catalogues. We regard this as the "reference" PDF since, as shown by Szapudi & Pan (2004) and checked by us, when the sampling is dense, all the above reconstruction methods recover the PDF of the mass, $P(N_g)$ and the mean biasing function very accurately. In the plot we show $P(1+\delta_g)(1+\delta_g)$ to highlight the low- and high-density tails, where the reconstruction is more challenging. The reconstructed PDF underestimates the reference PDF in the low- and high-density tails and overestimates it at $\delta \sim 0$. Systematic deviations in the low- and high-density tails are to be expected since the probability of finding halos, and therefore mock galaxies, in these regimes significantly deviates from the probability expected for a Poisson distribution. However, these differences are well within the 1-σ uncertainty strip as shown in the bottom panels of each plot.

The Γ_E method used to reconstruct the galaxy PDF from discrete counts is implemented as follows:

– We consider as the input dataset one of the volume-limited, luminosity complete subsamples listed in Table 1. The position of each object in the catalogue is specified in redshift space, i.e. by its angular position and measured spectroscopic redshift.
– Spherical cells are thrown at random positions within the surveyed region. We consider cells with radii $R = 4$, 6, and 8 h^{-1} Mpc . The smallest radius is set to guarantee $\langle N_g \rangle \geq 0.3$. The largest radius is set to have enough cell statistics to sample $P(N_g)$ at large N_g. We only consider cells that overlap by more than 60 % with the observed areas. This constraint reduces deviations from Poisson statistics (Bel & Marinoni 2014). Counts in the partially overlapping cells are weighted by the fraction f of the surveyed volume in the cell: N_g/f. The probability function $P(N_g)$ is then computed from the counts frequency distribution.
– We use the measured $P(N_g)$ and its moments to model the galaxy PDF with the Γ_E method that we compute using all factorial moments up to the sixth order.

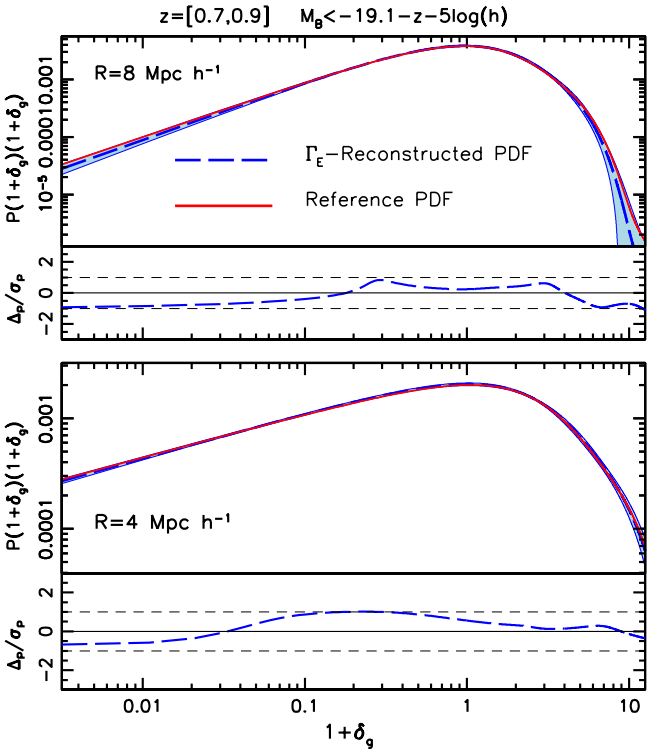

Fig. 2. Reconstructed PDF of the mock VIPERS galaxies measured in cells of $R = 4\ h^{-1}$ Mpc (top) and $R = 8\ h^{-1}$ Mpc (bottom). The blue solid curve represents the reference galaxy PDF obtained by averaging over the PDFs reconstructed from the *Parent* mocks using the Γ_E method. The blue dashed curve shows the average PDF reconstructed from the *Realistic* mocks using the same method. The blue shaded region represents the 1-σ scatter among the 26 *Realistic* mocks. We plot $P(1+\delta_g)(1+\delta_g)$ to highlight the performance of the reconstruction at high and low over-densities. We note the different Y-ranges in the two panels. The bottom panels in each plot show the difference Δ_p between the reconstructed and reference PDFs in units of the random error σ_p. Horizontal, dashed lines indicate systematic errors equal to 1-σ_p random uncertainties.

3.4. ....and from $P(\delta_g)$ to $b(\delta)\delta$.

To estimate the mean biasing function from the galaxy PDF, we solve Equation 3. To do so, we assume that shot noise is the main source of stochasticity and that a reliable model for the mass PDF is available. Despite its conceptual simplicity, this procedure requires several non-trivial steps that we describe below. The uncertainties introduced in each step are estimated in the next section. The procedure is as follows:

– We start from the galaxy PDF estimated from the measured $P(N_g)$, as described in the previous section.
– We assume a model PDF for the mass density field in redshift space. Rather than adopting some approximated, analytic model, we measure the mass PDF directly from a dark matter only N-body simulation with the same characteristics and cosmological model as the Millennium run (Springel

et al. 2005), thatis not based on the same model used to build the HOD-mock VIPERS catalogues. The use of an incorrect mass PDF is yet another possible source of systematic errors that we quantify in Section 4. However, this error is expected to be small since $\hat{b}$ and $\tilde{b}$ are mainly sensitive to σ and their ratio is largely independent of the underlying cosmology (Sigad et al. 2000).

- After computing the cumulative distribution function from the mass and galaxy PDFs, we use Eq. 3 to estimate the mean biasing function.
- We determine the maximum over-density $\delta_{\rm MAX}$ at which the reconstructed mean biasing function can be considered reliable. To estimate $\delta_{\rm MAX}$ we compare the measured $P(N_g)$ with the estimated $P(N_g)$ following the procedure described in Section 4.4.4 .
- We estimate the second-order moments $\hat{b}$ and $\tilde{b}$ and their ratio by integrating over all δ up to $\delta_{\rm MAX}$

$$\begin{aligned} \hat{b} &= \sigma^{-2} \int_{-1}^{\delta_{\rm MAX}} b(\delta)\delta^2 P(\delta)\, d\delta\,, \\ \tilde{b}^2 &= \sigma^{-2} \int_{-1}^{\delta_{\rm MAX}} (b(\delta)\delta)^2 P(\delta)\, d\delta\,. \end{aligned} \quad (7)$$

and test the robustness of the result with respect to the choice of $\delta_{\rm MAX}$.

4. Error sources

In this Section we review all possible sources of uncertainty that might affect the recovery of the biasing function and assess their amplitude using mock catalogues. In this process we need to consider a reference biasing function to compare with the results of the reconstruction. This could be estimated directly from the distribution of the dark matter particles and mock galaxies within the simulation box. However, we use the mean biasing function obtained from the *Parent* mocks as reference. We justify this choice as follows. First, Szapudi & Pan (2004) showed that when the sampling is dense both the Richardson-Lucy and the skewed lognormal fit methods recover the mean biasing function with high accuracy. Second, in Section 3.3 we found that when the sampling is dense the Γ_E method accurately recovers the mean biasing function in the *Parent* mocks.

4.1. Sensitivity to the galaxy PDF reconstruction method

Most of the previous estimates of the mean biasing function did not attempt to account for shot noise directly. This choice can hamper the recovery of $b(\delta)\delta$ when the sampling is sparse. To estimate errors induced by ignoring shot noise and quantify the benefit of using the Γ_E method we compared the biasing functions reconstructed using both procedures. The result of this test is shown in Figure 3. The red curve represents the reference biasing function obtained by averaging over the *Parent* mocks. In each mock the biasing function was estimated from the galaxy PDF using the Γ_E method. The blue dashed curve represents the same quantity estimated from the 26 *Realistic* mocks using the Γ_E method. The blue band represents the 2-σ scatter. For negative values of δ_g the reconstructed biasing function is below the reference biasing function, but the trend is reversed for $\delta_g > 0$, reflecting the mismatch between the reconstructed and reference PDFs in Figure 2. The discrepancy however, is mostly within the 2-σ scatter (horizontal dashed line in the bottom sub-panels). On the contrary, the biasing function obtained from the "direct" estimate of δ_g (brown dot-dashed curve and the corresponding 2-σ scatter, orange band) is significantly different from the reference function. The discrepancy increases at low densities and for small spheres, i.e. when the counts per cell decrease and the shot noise is large.

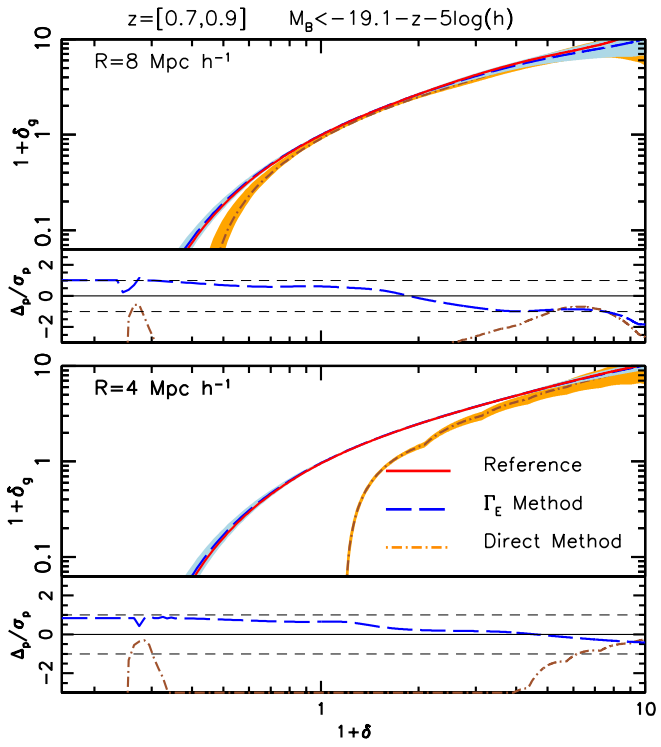

Fig. 3. Mean biasing function of mock VIPERS galaxies computed from counts in cells of $R = 4\ h^{-1}$ Mpc (bottom panel) and $R = 8\ h^{-1}$ Mpc (top panel). The magnitude cut and redshift range of the mock VIPERS subsample, indicated in the plot, are the same as Figure 2. Solid red curve: reference biasing function obtained from the *Parent* mock catalogues. Blue dashed curve and blue-shaded region: average value and 2-σ scatter of the biasing function reconstructed from the *Realistic* mocks using the Γ_E method.Brown dot-dashed curve and orange-shaded band: average value and 2-σ scatter of the biasing function reconstructed from the *Realistic* mocks using a 'direct' estimate of the galaxy PDF. Bottom sub-panels: difference Δ_p between the reconstructed and reference PDFs in units of the random error σ_p. Dashed lines indicate systematic errors equal to 1-σ_p random errors.

4.2. Sensitivity to the mass PDF

Another key ingredient of the mass reconstruction is the mass PDF. In principle this quantity could be obtained from galaxy peculiar velocities or gravitational lensing. However, in practice, errors are large and would need to be averaged out over scales much larger than the size of the cells considered here. For this reason we need to rely on theoretical modelling. Coles & Jones (1991) and Kofman et al. (1994) found that the mass PDF can be approximated by a lognormal distribution and this model was indeed adopted in previous reconstructions of the biasing function (e.g. M05, Wild et al. (2005), K11).

However, the lognormal approximation is known to perform poorly in the high- and low-density tails and for certain spectra of density fluctuations. An improvement over the lognormal model is represented by the skewed lognormal distribution (Colombi 1994). This model proved to be an excellent approx-

imation to the PDF of the dark matter measured from N-body experiments over a wide range of scales and of over-densities (Ueda & Yokoyama 1996). The impact of adopting either model for the mass PDF can be appreciated in Fig. 4. The solid red curves represent the same biasing functions shown in Fig. 3 obtained from the galaxy PDFs of the *Parent* mocks and from a mass PDF obtained directly from an N-body simulation with the same cosmological parameter and size as the Millennium simulation using the output corresponding to $z = 0.8$. As in the previous test, we consider the red solid curve as the reference biasing function. The brown dot-dashed curve shows the mean biasing function reconstructed assuming a lognormal model for the mass PDF, i.e. a lognormal fit to the PDF measured from the N-body simulation. The curve represents the average among 26 mocks and the orange band is the 2-σ scatter. For $R = 8$ h^{-1} Mpc , the biasing function is systematically below the reference whereas for $R = 4$ h^{-1} Mpc is above the reference at both high and low densities. The mismatch is very large and significantly exceeds the 1-σ scatter (bottom sub-panels). The skewed lognormal model (blue dashed curve) performs significantly better with differences well below 1-σ except at very negative δ values.

We conclude that, for the practical purpose of reconstructing galaxy bias, the mass PDF measured from N-body data and a skewed lognormal fit perform equally well. The main advantage of using the latter would be the possibility of determining the four parameters of the fit experimentally. Since, however, the parameters are poorly constrained by observations, we decided to adopt the mass PDFs from N-body simulations. This choice introduces a dependence on the cosmological model, however, that is mostly captured by one single parameter, σ, for which $\hat{b}$ and $\tilde{b}$ exhibit a linear dependent. With respect to this, the mass PDF used to obtain the biasing functions in Figure 4 is not the *true* mass PDF since it is obtained from an N-body simulation that uses a cosmological model that is different from the model used to produce the mock catalogues. We did this on purpose to mimic the case of the real analysis for which the underlying cosmological model is not known.

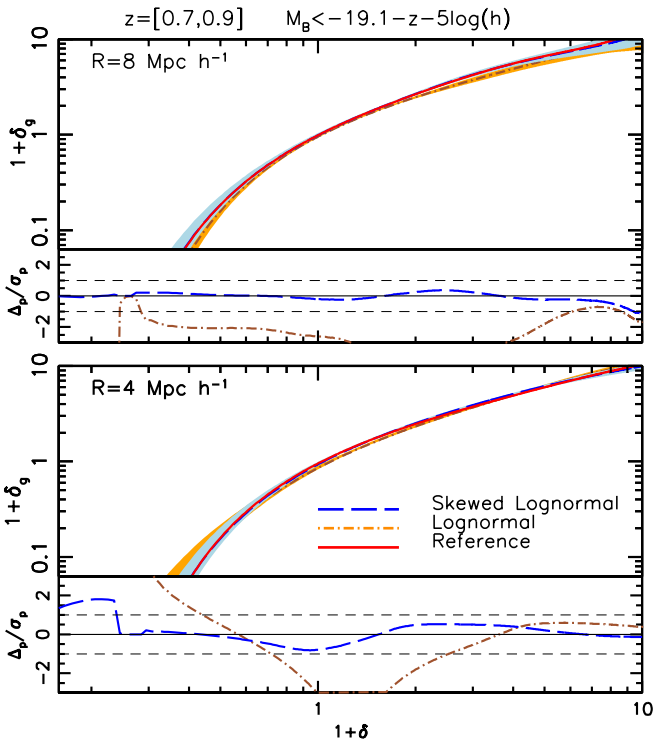

Fig. 4. Solid red curve: reference mean biasing function of Fig. 3 computed using the mass PDF from N-body simulations. Brown dot-dashed curve and orange band: biasing function obtained using a lognormal fit to the mass PDF and 2-σ scatter from the mocks. Blue dashed curve and blue band: biasing function obtained using a skewed lognormal fit to the mass PDF and 2-σ scatter from the mocks. Bottom panels: difference Δ_p between the reconstructed and reference PDFs in units of the random error σ_p. Dashed lines indicate systematic errors equals to 1-σ_p random errors.

4.3. Sensitivity to redshift distortions

Galaxy positions are measured in redshift space, i.e. using the observed redshift to estimate the distance of the objects. The presence of peculiar velocities induces apparent radial anisotropies in the spatial distribution of galaxies and, as a consequence, modifies the local density estimate and their PDF (Kaiser 1987). However, our goal is to reconstruct the mean biasing function in real space without redshift distortions. Considering the difficulties and uncertainties in determining the galaxy PDF in real space, one could instead consider the galaxy and mass PDFs both measured in redshift space under the assumption that peculiar velocities induce similar distortions in the spatial distribution of both dark matter and galaxies so that they cancel out when estimating the mean biasing relation from Eq. 3. In the limit of the Gaussian field, linear perturbation theory and no velocity bias, the cancelation is exact. However, non-linear effects have a different impact on the mass and galaxy density fields and induce different distortions in their respective PDFs. To assess the impact of these effects we compared the mean biasing function of mock galaxies reconstructed from PDFs estimated in real and redshift space.

The results are shown in Figure 5. The solid red curve represents the mean biasing function of galaxies in the *Realistic* mock catalogues estimated using the PDFs of galaxies and mass in real space. The blue dashed line shows the same function estimated in redshift space. Both curves are obtained by averaging over the 26 mocks and the blue band represents the 2-σ scatter in redshift space. The redshift space biasing function underestimates the true biasing function in low-density regions and overestimates it at high densities, i.e. in the presence of highly non-linear flows. The difference is systematic but its amplitude is within the 2-σ random errors estimated by adding in quadrature the scatter among mocks in real and redshift space (bottom panels in each plot). The biasing functions shown in Figure 5 represents a demanding test in which we consider the smallest cells of 4 h^{-1} Mpc where deviations from linear motions are larger. The discrepancy decreases if the size of the cell increases.

These systematic differences induce errors in the estimated moments $\hat{b}$ and $\tilde{b}$. To quantify the effect we computed the moments as a function of δ (i.e. by varying $\delta_{\rm MAX}$ in Eq. 7) both in real and redshift space. The results are shown in Figure 6. The plots show the per cent difference between the moments measured in redshift versus real space. The panels and curves refer to the same redshift bins and magnitude cuts as in Figure 5. Systematic errors induced by redshift distortions are ~ 2 % for $\hat{b}$ and for $\tilde{b}$ (not shown) and one order of magnitude smaller for $\tilde{b}/\hat{b}$. They provide the main contribution to the total systematic errors listed in Table 2 and are of the same size, although somewhat smaller than the random errors.

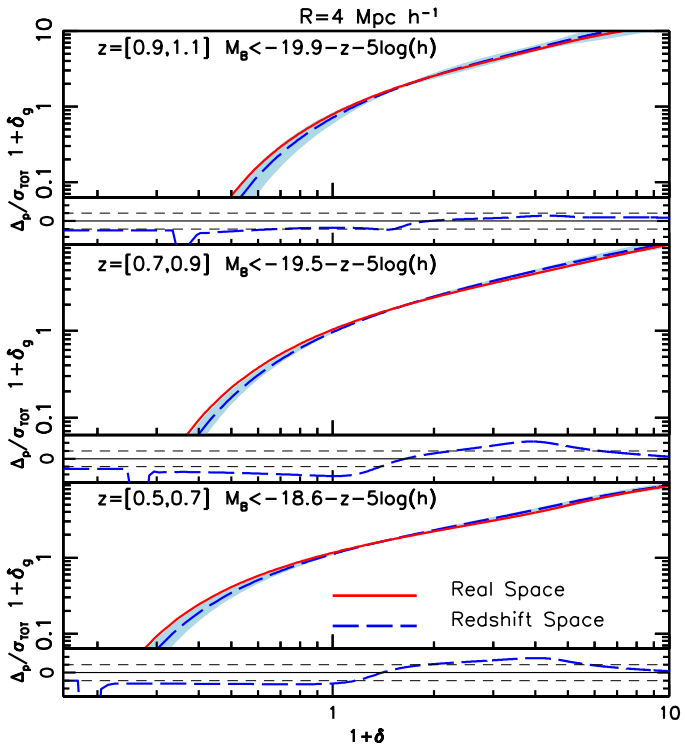

Fig. 5. Mean biasing function estimated in real space (solid, red curve) and redshift space (dashed blue curve and its 2-σ uncertainty band). Counts are performed in spherical cells with a radius of 4 h^{-1} Mpc . The luminosity cut and the redshift range is indicated in each panel. The width of each band represents the scatter among mocks. In the bottom part of each plot we show the difference Δ_p between the reconstructed and reference PDFs in units of $\sigma_{\rm TOT}$, where $\sigma_{\rm TOT}$ accounts for the *rms* scatter in both the real- and redshift-space mocks. Dashed lines indicate where systematic errors equal to 1-$\sigma_{\rm TOT}$ random errors.

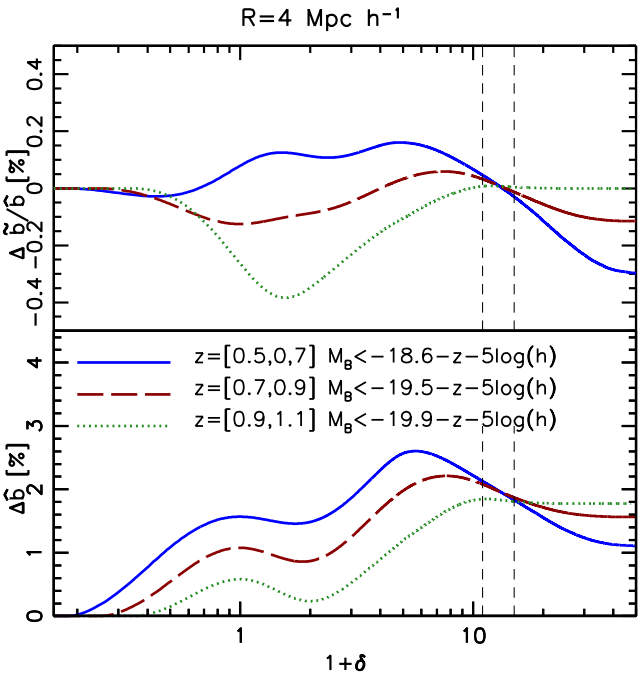

Fig. 6. Bottom panel: per cent difference between the $\hat{b}$ values estimated in redshift space and in real space using spherical cells with a radius of 4 h^{-1} Mpc as a function of $1+\delta_{\rm MAX}$ (see Eq. 7). The different curves refer to different redshift shells and magnitude cuts, as indicated in the plot. Upper panel: per cent difference in the estimated non-linear parameter $\tilde{b}/\hat{b}$. Vertical dashed lines are drawn in correspondence to the $\delta_{\rm MAX}$ values at which systematic errors are computed and listed in Table 2.

Considering the absolute and relative size of these errors, we perform our analysis in redshift space.

4.4. Error estimate

Different sources of errors affect the recovery of the biasing function. One error source is cosmic variance due to the finite volume of the sample. This source dominates the error budget of the M05 and K11 analyses.

The other sources are the shot noise induced by discrete sampling and the limited number of independent cells used to build the probability of galaxy counts $P(N_g)$. In the VIPERS survey, which is based on a single-pass strategy, sparse sampling is more of an issue than in the M05 and K11 cases. The cumulative effect of the single pass strategy and colour preselection reduces the sampling rate to $\sim$ 35 % on average with significant variations across quadrants. The survey geometry, characterised by gaps and missing quadrants that occupy $\sim$ 25 % of the would-be continuous field, further dilutes the sampling (we consider cells that overlap up to 40 % with unobserved regions) and limits the number of independent cells that can be accommodated within the survey. Our PDF reconstruction strategy is designed to minimise these effects that, nevertheless, induce random and systematic errors that need to be estimated. We do this with the help of both the *Parent* and *Realistic* mock catalogues. The former provide the reference mean biasing function. Errors are estimated by comparing the bias function reconstructed from the *Realistic* mocks to the reference mocks. The procedure is detailed below and the estimated errors are listed in Table 2.

4.4.1. Total random error

To estimate the total random error $\sigma_{\rm RND}$, we proceed as follows. We reconstruct the mean biasing function in each of the *Realistic* mock catalogues, compute the average over the 26 mocks and, finally, estimate the scatter around the mean. The *rms* scatter provides an estimate of the total random error. All sources of uncertainties contribute to this error (e.g. cosmic variance, shot noise, and limited number of cells), which may affect the recovery of the biasing function. Total random errors for $\hat{b}$ and $\tilde{b}/\hat{b}$ are listed in columns 6 and 10 of Table 2, respectively.

4.4.2. Cosmic variance

To assess the contribution of the cosmic variance, $\sigma_{\rm CV}$, to the error budget, we proceed as for the estimate of total random errors using, however, the *Parent* catalogues rather than the *Realistic* catalogues. Since errors in the bias reconstruction are mainly driven by discrete sampling and in the *Parent* catalogues the sampling is dense, the *rms* scatter among these mocks is dominated by cosmic variance. Cosmic variance contributions to errors in $\hat{b}$ and $\tilde{b}/\hat{b}$ are shown in columns 7 and 11 of Table 2, respectively. It turns out that the contribution of the cosmic vari-

ance is of the same order as that of the sparse sampling and, unlike in the case of M05 and K11, it does not dominate the error budgets.

4.4.3. Systematic errors

Following K11, we compute systematic errors, $\sigma_{\rm SYS}$, as the average offset of the bias estimates in the *Realistic* and the *Parent* catalogues, i.e. $\sigma_{\rm SYS} = \langle X_{Realistic} - X_{Parent} \rangle$, where X is either $\hat{b}$ or $\tilde{b}/\hat{b}$ and the mean is over the 26 pairs of mocks. These systematic errors are plotted in the bottom panels of Fig. 3 (blue, dashed curves). Their amplitudes at $\delta_{\rm MAX}$ are listed in columns 8 and 12 of Table 2. These systematic errors are of the same order as the random errors and as the errors induced by redshift distortions discussed in Section 4.3. These systematic errors include those induced by redshift distortions. The fact that they are of the same order as those discussed in Section 4.3 indicates that they dominate the budget of systematic errors.

Our systematic errors are similar to those estimated by K11 (upper part of their Table 2) from the zCOSMOS sample, which is significantly small than VIPERS. As these errors do not seem to depend on the volume of the survey, we conclude that they can be regarded as genuinely systematic. Systematic errors on $\hat{b}$ are on average positive, meaning that the mean slope of the reconstructed biasing function typically overestimates the true biasing function. For the non-linear *bias,* parameter systematic errors are preferentially negative, indicating that the reconstruction procedure has the tendency to underestimate the non-linearity of the biasing function.

4.4.4. The value of $\delta_{\rm MAX}$

Our *bias* estimator becomes progressively less reliable as the density increases, for two reasons: first, the numerical solution to Eq. 3 becomes unstable when the cumulative distribution functions approach unity, i.e. in correspondence of the high peaks of the mass and galaxy density fields. In this regime, small errors in the estimated galaxy PDF propagate into large uncertainties in δ; second, as anticipated in the previous Section, the scatter in the δ_g versus δ relation is larger than Poisson. Our assumption that Eq. 6 is valid at all δ leads to underestimating the high-density tail of the galaxy PDF and, consequently, the value of $\hat{b}$.

Our mock catalogues can be used to estimate the first type of error, but cannot fully account for the second type of error since our mock galaxies are sampled from dark matter halos assuming Poisson statistics. We therefore take the alternative route of reducing the impact of deviations from Poisson statistics at high densities. We do this by setting a sensible maximum over density value, $\delta_{\rm MAX}$, at which we compute the bias moments. The value of this threshold is computed as follows:

1. We consider the difference ΔP between the 'true' $P^t(N_g)$ measured in the *Realistic* mock catalogues and the reconstructed $P^r(N_g)$ estimated through Eq. 5.
2. We search for the first N_g value, N_1, at which $\Delta P > 2\sigma_P$, where σ_P is the *rms* scatter in the mocks.
3. We search for the first N_g value, N_2, at which $|\Delta P/P^t(N_g)| > 0.5$.
4. We take $N_{\rm MAX} = Min[N_1, N_2]$ and compute the corresponding over-density in galaxy counts $\delta_{g,\rm MAX} = N_{\rm MAX}/\langle N \rangle$.
5. We obtain the corresponding mass over-density $\delta_{\rm MAX}$ from $\delta_{g,\rm MAX}$ from the estimated mean biasing function.

The largest over-density at which we search for a solution to Eq. 3 is $\delta_{\rm MAX}$, and this is also the over-density at which we estimate the bias moments. This value is clearly model dependent since it was estimated from the VIPERS mocks. An alternative way of setting this threshold would be to look for wiggles in the mean biasing function measured from real data, i.e. spurious features induced by instabilities in the reconstruction procedure. We found that this second criterion is less stringent as it produces $\delta_{\rm MAX}$ values larger than using mocks. We decided to adopt a conservative approach and use the $\delta_{\rm MAX}$ thresholds estimated with the first procedure.

With this criterion we obtain different $\delta_{\rm MAX}$ for the different galaxy subsamples considered in our analysis. This limits our ability to compare results. Since the value of $\delta_{\rm MAX}$ mainly depends on the radius of the cell, we use one single value for $\delta_{\rm MAX}$ for a given cell size, irrespective of the other parameters used to define the subsample. These values, which are listed in Table 2, correspond to the minimum $\delta_{\rm MAX}$ among those computed for all subsamples.

All bias parameters presented in our work were computed at these over-density values. To check the robustness of our results to $\delta_{\rm MAX}$ we also considered a second, less stringent threshold obtained by taking the maximum value of $\delta_{\rm MAX}$ among those of the various subsamples for a given cell size. This second set of $\delta_{\rm MAX}$ that we denote as $\bar{\delta}_{\rm MAX}$, is also listed in Table 2 together with the corresponding estimates for the bias moments (values in parenthesis).

5. Results

In this section we present the results of our analysis, focusing on the dependence of the mean biasing function and its moments on various quantities. In Sections 5.1 and 5.2 we explore the bias dependence on magnitude and redshifts, respectively. In both cases we fix the radius of the cells equal to 6 h^{-1} Mpc . The dependence on the cell size is investigated in Section 5.3. Results are summarised in Section 5.4 and listed in Table 2.

5.1. Magnitude dependence

The different solid curves in Figure 7 represent the mean biasing function of VIPERS galaxies reconstructed from counts in cells of radius 6 h^{-1} Mpc for different magnitude cuts for three different redshift shells (the three panels). We applied a small horizontal offset $\delta = 0.015$ to the curves to avoid overlapping error bars. We plot $(1 + \delta)$ in logarithmic units both to ease the comparison with similar plots in the literature and to highlight deviations from linearity in the low-density regions. Error bars represent the 2-σ random scatter computed from the *Realistic* mocks.

The magnitude range that we are able to explore is set by competing constraints: the faint limit reflects the requirement of maximising the completeness of the sample whereas the bright limit is set by requiring $\langle N_g \rangle > 0.3$ per cell. As a result, the magnitude range shrinks with the redshift: at $z = [0.5, 0.7]$ it spans a range $\Delta M_{\rm B} = 1.4$ whereas at $z = [0.9, 1.1]$ $\Delta M_{\rm B} = 0.5$.

In the upper plot the curves corresponding to the different magnitude cuts are well separated for $\delta_g < 0$. The separation reduces and then disappears with the redshift. This is not surprising since at $z \geq 0.9$ the luminosity range is very narrow, as we have seen. No significant trend with luminosity is seen at large over-density. These features, or the lack of them, are robust to variations in the size of the cells in the range $R = [4, 8]$

h^{-1} Mpc (see Table 2) and confirm the results obtained at lower redshifts from galaxy clustering (e.g. Norberg et al. 2002; Zehavi et al. 2005; Pollo et al. 2006; Coil et al. 2008; Skibba et al. 2014; Arnalte-Mur et al. 2014; Marulli et al. 2013), gravitational lensing (e.g. Coupon et al. (2012)) and counts in cells (e.g. M05 and K11).

To further investigate galaxy bias in under-dense regions, we zoom into the $\delta < 0$ range in Figure 8. The curves are the same as in Figure 7. The black long-dashed line represents the linear biasing function with a slope matching the $\hat{b}$ value estimated at δ_{MAX}, which is listed in Table 2. Since $\hat{b}$ only weakly depends on the magnitude cut we only consider one representative case per panel. The local slope of the biasing function is always steeper than the best-fitting linear bias model. The horizontal, short-dashed line shows the $\delta_g = -0.9$ threshold.The mass over-density at which this line crosses the biasing curves, $\delta_{\rm TH}$, increases with the redshift and, to a lesser extent, with the luminosity. This trend, which was noticed by M05 and, with less significance, by K11,has been interpreted as evidence that low-density regions are preferentially populated by low-luminosity galaxies. Also, the quantity $\delta_{\rm TH}$ has been regarded as the typical mass over-density below which very few galaxies form.

Figure 8 shows that galaxies can be found at mass over-densities well below $\delta_{\rm TH}$. This low-density tail, together with the steepness of the biasing function for $\delta > \delta_{\rm TH}$, shows that the biasing relation in the under-density region significantly deviates from the linear prescription. Non-linearity increases when decreasing the cell size. As we checked, for $R = 4\ h^{-1}$ Mpc the slope of the biasing curves further increases well above $\delta_{\rm TH}$. For $R = 8\ h^{-1}$ Mpc , the difference disappears and the two slopes start to match. Still, the *bias* curves keep featuring a negative δ tail that cannot be matched by linear models.

Figure 9 shows the second-order moment $\hat{b}$ (left panels) and the ratio $\tilde{b}/\hat{b}$ (right panels) of the biasing functions shown in Fig. 7. The same colour-code is used to indicate the magnitude cuts. Large filled symbols refer to measurements performed at $\delta_{\rm MAX}$ assuming $\sigma_8 = 0.9$ whereas the slightly offset, smaller open symbols refer to estimates performed at $\bar{\delta}_{\rm MAX}$. The values of the corresponding bias moments are listed in Table 2. Error bars represent 1-σ total random uncertainties estimated from the *Realistic* mocks (see Table 2).

In the left panels of Fig. 7, we notice that in the low redshift bin, where the magnitude interval that we probe is larger, $\hat{b}$ increases with the luminosity. This dependence is much weaker for $z = [0.7, 0.9]$ and completely absent at higher redshifts. We show results for cells of 6 h^{-1} Mpc . However, the same trend is also seen for 4 and 8 h^{-1} Mpc .

The right panels show the non-linear parameter $\tilde{b}/\hat{b}$. Values that differ from unity indicate deviations from linear bias (horizontal dashed line). A small but significant degree of non-linearity is present at all redshifts. We do not detect any significant dependence on luminosity in any redshift bin and for any cell size.

A common feature of the reconstructed mean biasing functions at $z = [0.9, 1.1]$ is the presence of some irregular behaviour (wiggling) at high over-densities. This is the typical fingerprint of an imperfect inversion (Eq. 3) discussed in Section 3.2 and one of the reasons for introducing the threshold $\delta_{\rm MAX}$. These irregularities typically arise as a result of sampling rare, large over-densities with a limited number of independent cells. The effect is most evident at large redshifts and for bright magnitude cuts, i.e. when the sampling is sparser. This affects the shape of the reconstructed mean biasing function. However, the impact on

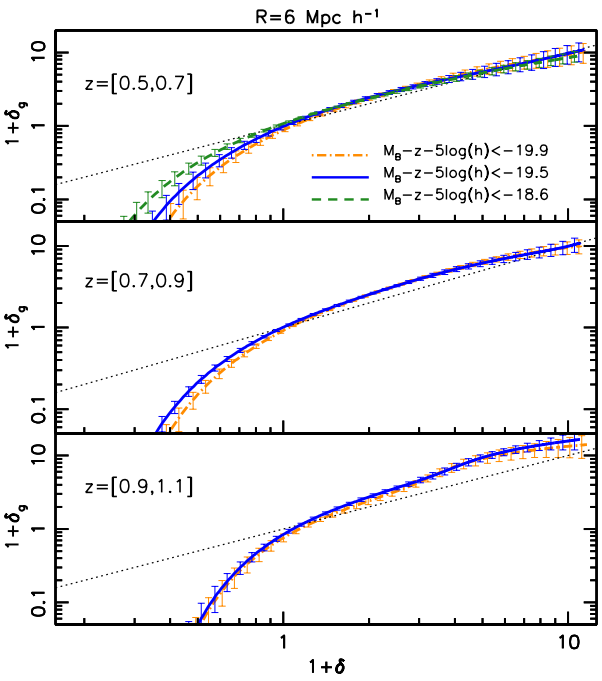

Fig. 7. Mean biasing function of VIPERS galaxies from counts in cells of radius 6 h^{-1} Mpc as a function of the B-band magnitude cut in three redshift ranges indicated in each panel. Curves with different colours and line styles correspond to the different magnitude cuts indicated in upper panel. Error bars with matching colours represent the associated 1-σ uncertainty intervals estimated from the mocks. A horizontal offset $\delta = 0.015$ was applied to avoid overlapping error bars. All biasing functions are plotted out to $\delta_{\rm MAX}$.

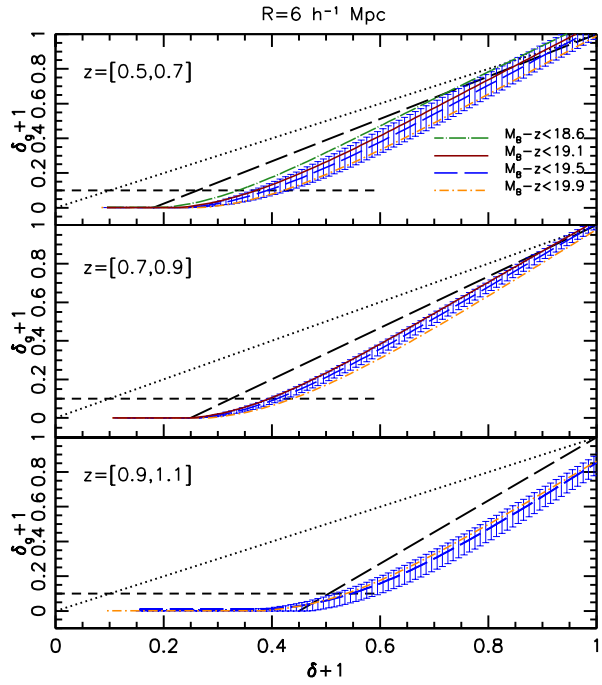

Fig. 8. Zoom into the under-density range of Fig. 7. The horizontal short-dashed line represents the over-density threshold $\delta_g = -0.9$. The long dashed line shows the linear biasing function $\delta_g = \hat{b}\delta$, for the $\hat{b}$ value corresponding to the $M_{\rm B} - z - 5log(h) < 19.5$ cut, which is listed in Table 2. The dotted line, shown for reference, shows the case $b_{LIN} = 1$. Error bars represent the 1-σ *rms* scatter among the mocks. They are only shown for the case $M_{\rm B} - z - 5log(h) < 19.5$ to avoid overcrowding.

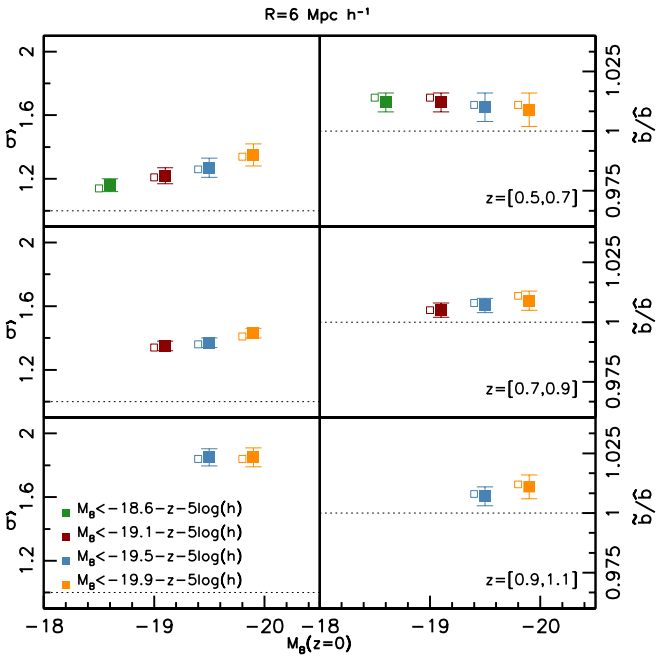

Fig. 9. Second-order moments of the mean biasing functions shown in Fig. 7. Left panels: moment $\hat{b}$. Right panels: non-linear parameter $\tilde{b}/\hat{b}$. The cell size is $R = 6\ h^{-1}$ Mpc . Error bars indicate 1-σ scatter from the mocks. The redshift ranges and colour code are the same as in Fig. 7. Magnitude cuts are indicated in the plots. All values were computed assuming $\sigma_8 = 0.9$. The horizontal dashed line is plotted for reference and represents the case of no bias (left plot) and linear bias (right panels). Large filled symbols refer to measurements performed at $\delta_{\rm MAX}$, and small, open symbols refer to estimates at $\bar{\delta}_{\rm MAX}$.

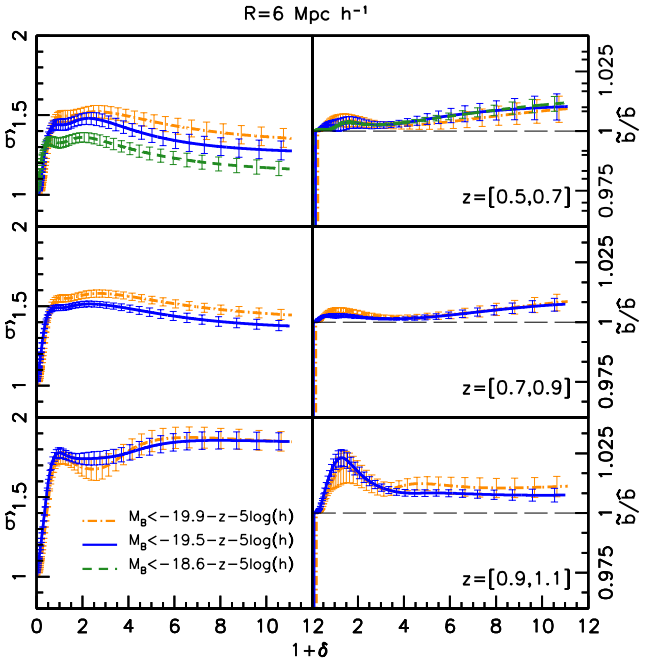

Fig. 10. Left: second-order moment $\hat{b}(\delta)$ of the reconstructed mean biasing functions shown in Fig 7. The cell size is $R = 6\ h^{-1}$ Mpc . Different line styles and colours indicate different luminosity cuts listed in the plot. The redshift ranges and colour codes are the same as in Fig 7. Error bars represent the *1*-σ scatter among the mocks. Right: similar plots showing the non-linear bias parameter $\tilde{b}/\hat{b}(\delta)$. A horizontal offset $\delta = 0.015$ was applied to avoid overlapping error bars. All curves are plotted out to $\delta_{\rm MAX}$.

the second moments $\hat{b}$ and $\tilde{b}$ and, especially, $\tilde{b}/\hat{b}$, is rather limited. This is because bias moments are integral quantities (Equation 7) weighted by the mass PDF, which peaks at $\delta \sim 0$ and rapidly approaches zero in the high- and low-density tails. Systematic errors in the bias reconstruction at large over-densities are therefore suppressed when computing $\hat{b}$ and $\tilde{b}$ and further smoothed out when computing their ratio.

Figure 10 demonstrates the validity of this conjecture. In the left panels we show the values of $\hat{b}(\delta)$ computed from equation 7. Curves with different line styles refer to the different magnitude cuts indicated in the plot. Error bars with matching colours indicate the 1-σ scatter from the mocks. In the interval $z = [0.9, 1.1]$ and for the brightest and sparsest sample, $\hat{b}(\delta)$ flattens for $\delta > 3$, i.e. well below $\delta_{\rm MAX}$. Analogous considerations hold for the curve $\tilde{b}/\hat{b}(\delta)$ shown in the right panels. These trends are robust to the size of the cells.

5.2. Redshift dependence

To explore the bias dependence on the redshift we set a magnitude cut $\mathrm{M_B} = -19.5 - z - 5\log(h)$ and estimated the mean biasing function of brighter galaxies in the three redshift bins. This z-dependent magnitude cut is designed to account for luminosity evolution (Zucca et al. 2009), so that differences in the galaxy bias measured in the different z-bins can be interpreted as the result of a genuine evolution. The results of our analysis are shown in Figure 11. The plots are analogous to those of Fig. 7 and use the same symbols, colour scheme, and line style. However, we consider cells of different sizes in the three panels.

The biasing function shows little or no evolution in the range $z = [0.5, 0.9]$, as demonstrated by the proximity between the dashed-blue ($z = [0.5, 0.7]$) and dot-dashed orange ($z = [0.7, 0.9]$) curves and the overlap of their 1-σ error bars. The red solid line, however, is separated fullyfrom the others, indicating that galaxy bias evolves significantly beyond $z = 0.9$. This evolution is detected both in low- and high-density environments. It implies that $\delta_{\rm TH}$ increases significantly with the redshift, indicating that evolution shifts galaxy formation towards regions of progressively lower density. At $\delta > 0$ the effect of evolution is that of increasing the slope of the biasing function with z. Since in this range the biasing is close to linear, an estimate of b_{LIN} would reveal a redshift evolution consistent with that observed in several analyses, as detailed in Section 6. The same trend is evident in all panels, indicating that the bias evolution is similar in all explored scales.

At high redshifts and for $R = 8\ h^{-1}$ Mpc (bottom panel) the biasing function is characterised by some irregularities at moderate values of δ. As pointed out, these have little impact on the estimated values of $\hat{b}$ and $\tilde{b}/\hat{b}$. It is reassuring that these anomalies are only seen at high redshifts, confirming the fact that they are induced by poor sampling of the counts probability. All this makes us confident that bias evolution is a genuine feature.

Figure 12 shows the values of $\hat{b}$ and $\tilde{b}/\hat{b}$ as a function of redshift. The values were obtained by integrating the mean biasing functions in Figure 11 out to the values $\delta_{\rm MAX}$ and $\bar{\delta}_{\rm MAX}$ (large and small symbols, respectively). The corresponding values are listed in Table 2. The colour code is the same as in Figure 11 and is indicated in the plot. The mean slope of the curve, $\hat{b}$ (left panels), increases significantly beyond $z = 0.9$ whereas we see

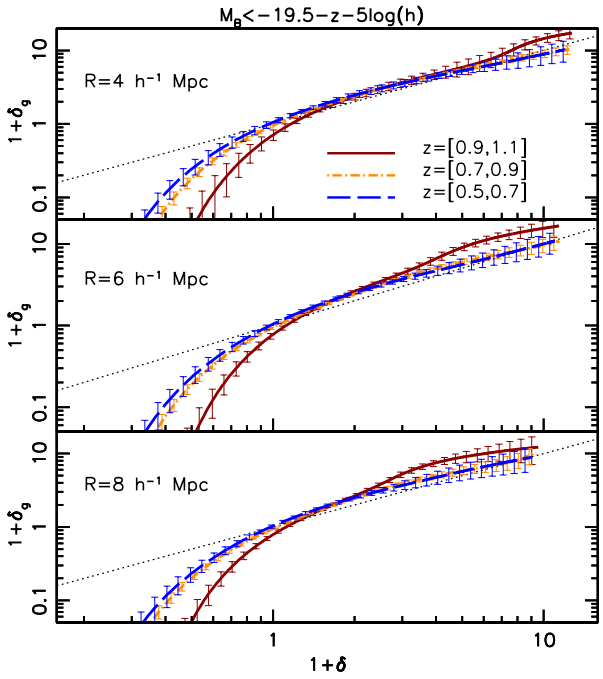

Fig. 11. Mean biasing function of VIPERS galaxies with $M_B < -19.5 - z - 5\log(h)$ measured in different redshift bins, characterised by different colours and line-styles, as indicated in the plot. Error bars represents the 1-σ scatter in the mocks. The three panels refer to different cell sizes with radii $R = 4, 6, 8\ h^{-1}$ Mpc from top to bottom. A horizontal offset $\delta = 0.015$ was applied to avoid overlapping error bars. All curves are plotted out to δ_{MAX}.

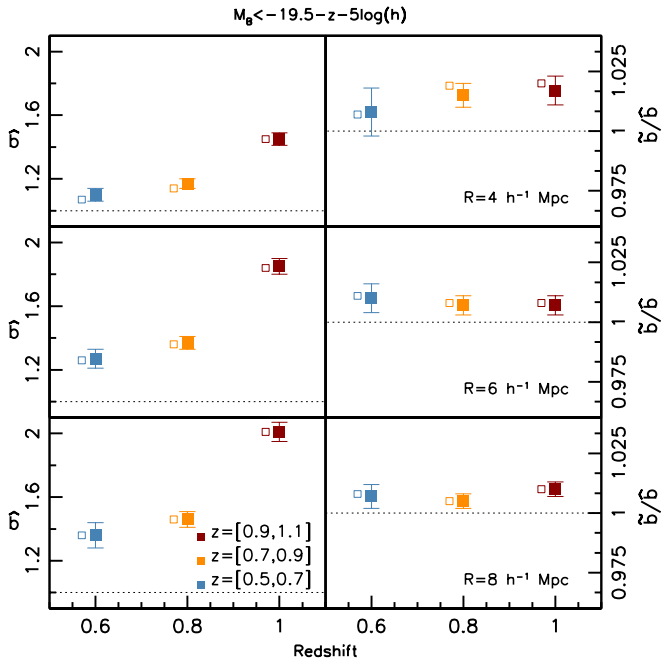

Fig. 12. Moments of the mean biasing functions shown in Fig. 11. Panels from top to bottom indicate cells of increasing size. Left panels: $\hat{b}$. Right panels: $\tilde{b}/\hat{b}$. Error bars indicate the 1-σ scatter from the mocks. All values were computed assuming $\sigma_8 = 0.9$. Large filled symbols refer to measurements performed at δ_{MAX}, small open symbols refer to estimates at $\bar{\delta}_{MAX}$.

little or no evolution at lower redshifts. This shows that the trend seen in Figure 11 is seen at all scales, indicating that the bias evolution at $z > 0.9$ is indeed a robust feature. The bias parameter of VIPERS galaxies brighter than $M_B = -19.9 - z - 5\log(h)$ exhibits a small but significant degree of non-linearity at all redshifts and scales explored in our analysis (right panels); this bias parameter, however, does not significantly evolve with redshift. These results are robust to the luminosity cut since they are also found for galaxies brighter than $M_B < -19.9 - z - 5\log(h)$.

5.3. Scale dependence

In Fig. 13 we explore the dependence of the bias of VIPERS galaxies on the radius of the cells down to a scale of 4 h^{-1} Mpc. In the plots we show the mean biasing function of VIPERS galaxies brighter than $M_B = -19.5 - z - 5\log(h)$ measured at $R = 4, 6$ and $8\ h^{-1}$ Mpc. Different scales are characterised by different colours, as indicated in the plot. The panels show the results in the three redshift shells. At negative over-density the curves are remarkably similar, indicating that δ_{TH} and the efficiency of galaxy formation do not depend on the scale in the range $[4, 8]\ h^{-1}$ Mpc. At $\delta > 0$ the curves steepen with the radius of the cell, indicating that biasing increases with the scale especially at high redshift.

A more quantitative assessment of scale dependence is shown in Fig. 14. The value of $\hat{b}$ steadily increases with the cell radius, R, especially at high redshift. This trend may sound counterintuitive; galaxies are expected to trace the mass with increasing accuracy on a larger scale and, consequently, galaxy bias is expected to approach its linear value. This, however, occurs on scales much larger than those considered here (see e.g.

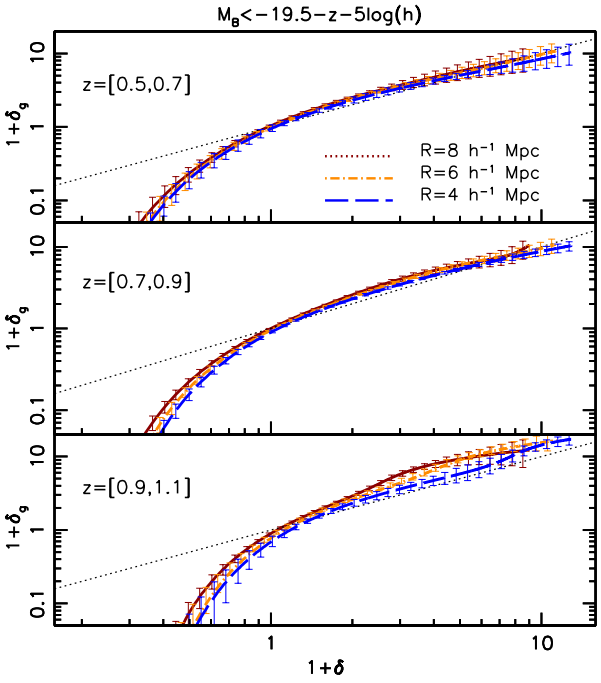

Fig. 13. Mean biasing function of VIPERS galaxies with $M_B < -19.5 - z - 5\log(h)$ computed from counts in cells with radii of 4, 6, and 8 h^{-1} Mpc. Biasing functions at different scales are indicated with different colours and line styles, as indicated in the plots. Error bars represent the 2-σ *rms* scatter in the mocks. Different panels refer to different redshift shells. An offset $\delta = 0.015$ was applied to avoid overlapping error bars. Curves are plotted out to δ_{MAX}.

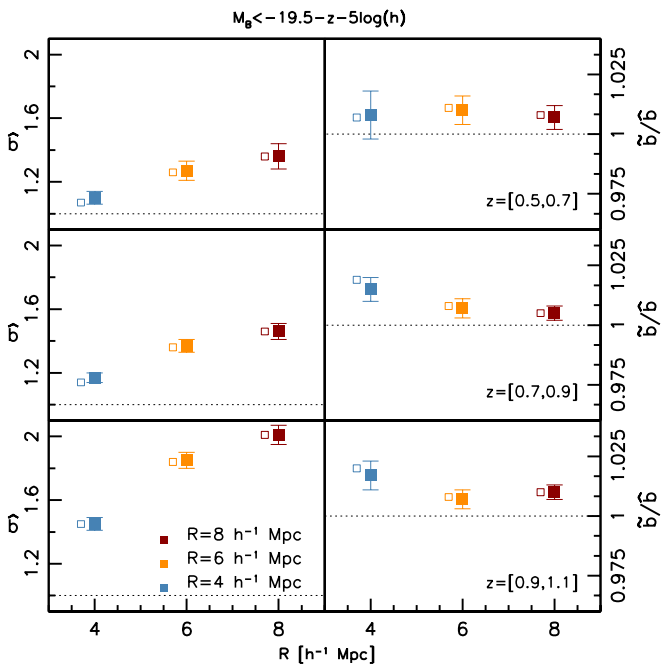

Fig. 14. Moments of the mean biasing functions vs. the size of the cell. Panels from top to bottom refer to different redshift ranges indicated in the plot. Left panels: $\hat{b}$. Right panels: $\tilde{b}/\hat{b}$. The parameters were computed assuming $\sigma_8 = 0.9$. Error bars indicate the 1-σ *rms* scatter from the mocks. Large, filled symbols refer to measurements performed at $\delta_{\rm MAX}$, small open symbols refer to estimates at $\bar{\delta}_{\rm MAX}$.

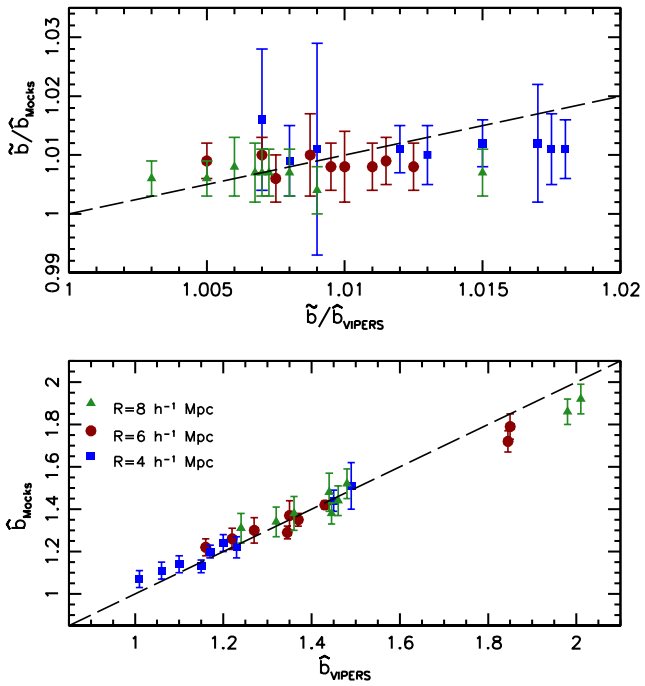

Fig. 15. Bottom panel: comparison between $\hat{b}$ measured in various mock subsamples and $\hat{b}$ measured in the VIPERS catalogue. Different symbols and colours refer to results obtained with cells of different sizes, as indicated in the plot. The subsamples were obtained by applying the same magnitude and redshift cuts used in this section and indicated in Table 2. The error bars represent 1-σ scatter from the mocks. Top panel: comparison among the non-linear parameters $\tilde{b}/\hat{b}$ measured in the mocks and in the real sub-catalogues
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Wild et al. 2005). On the scales explored here the halo model predicts that the opposite trend should be observed (see e.g. Fig. 4 of Zehavi et al. 2004). The reason is that in this range of scales the contribution to galaxy clustering of the 1-halo term, which dominates on small scales, is comparable to that of the 2-halo term, which dominates on large scales. The scale of the crossover depends on galaxy type and redshift but it is expected to be bracketed in the range probed by our analysis. This explanation is corroborated by the fact that the values of $\hat{b}$ measured in the HOD mocks, designed following the halo model prescriptions, do show an increasing trend with the size of the cells.

An increase of galaxy bias with the scale was already detected at lower redshifts from the analysis of galaxy clustering (Zehavi et al. 2005) and from weak lensing (Hoekstra et al. 2002; Simon et al. 2007). This is the first detection at relatively high redshift that exploits counts in cell statistics. A small, but significant amount of non-linearity is detected at all redshifts. Unlike $\hat{b}$, the non-linear parameter $\tilde{b}/\hat{b}$ seems to be scale independent. These results are robust to magnitude cut since similar trends for $\hat{b}$ and $\tilde{b}/\hat{b}$ are also seen when one restricts the biasing analysis to objects brighter than $\mathrm{M_B} < -19.9 - z - 5\log(h)$.

5.4. Results from the whole dataset

We now summarise the results presented in this section. Overall, the biasing functions of the VIPERS subsample are in qualitative agreement with those of M05 and K11 with some intriguing differences. At moderate over-densities and out to $\delta_{\rm MAX}$, our biasing functions are close to linear with a slope close to $\hat{b}(\delta_{\rm MAX})$. This is at variance with M05 and K11 whose biasing function flattens at large δ, leading to an anti-bias signature. This feature has been variously interpreted as evidence for quenching processes (Blanton et al. 2000), enhanced galaxy merging rate (Marinoni et al. 2005), and early galaxy formation (Yoshikawa et al. 2001) in high-density regions. We find a similar flattening only if we push our analysis beyond $\delta_{\rm MAX}$. However, its statistical significance is less that 2 σ. A similar feature was also detected in simulations and interpreted as an artefact due to limitations of the bias estimator at high redshift (Sigad et al. 2000). Given the fact that all these works, including ours, use a similar technique to measure galaxy bias, we suspect that the flattening at large density is not a genuine effect.

At $\delta < 0$ the biasing function significantly deviates from linear prescription and is steeper than $\hat{b}$. Moreover, the galaxy density remains positive below $\delta_{\rm TH}$, indicating that galaxy formation is not entirely quenched even in very low-density environments.

Table 2 lists the bias parameters measured in the VIPERS subsamples of Table 1 together with random and systematic errors estimated from the mocks. We computed all of the parameters by integrating the mean biasing function out to the value $\delta_{\rm MAX}$ listed in the Table. Altogether the results confirm the various trends that we described in the previous sections: the value of $\hat{b}$ increases with luminosity, scale, and with the redshift beyond $z = 0.9$. Deviations from linear biasing are small but typically detected with significance larger than 1 σ. The non-linear bias parameter is, within the errors, independent of redshift, luminosity, or scale.

We obtained errors from the VIPERS mock catalogues designed to match the 2-point statistics of real galaxies but not their abundance or their bias. One consequence of this is that

bright galaxies in the subsamples of the mocks are sparser than the real galaxies at $z < 0.9$ (see Table 1). As a result our random errors somewhat overestimate the real errors. As for the bias, if that of the galaxies in the real sample is different from the mock sample, then our error estimate would be affected. We compared the values of $\hat{b}$ and $\tilde{b}/\hat{b}$ in the mock and in the real samples to investigate this issue. The resulting scatter plots are shown in Figure 15. The different points represent the individual subsamples considered in our analysis. Symbols with different colours are used to highlight results obtained with different cell sizes. Error bars represent the *rms* scatter in the mocks. Most of the points deviates less than 2 σ from the expected value (black dashed line), implying that our mocks are realistic and that our errors are indeed reliable.

6. Comparison with previous results

Several authors estimated the bias of galaxies in the same range, $z = [0.5, 1.1]$, considered here. The majority of these authors assumed linear bias and estimated the bias parameter from galaxy clustering (Coil et al. 2006; Meneux et al. 2006; Coil et al. 2008; Meneux et al. 2008, 2009; Coupon et al. 2012; Marulli et al. 2013; Skibba et al. 2014; Arnalte-Mur et al. 2014). Only a handful of papers addressed the issue of non-linear or scale-dependent bias at these redshifts (M05, K11, Simon et al. (2007); Jullo et al. (2012)). In this section, we compare our results with both types of analyses. First we compare our estimated non-linear bias parameter with available measurements from previous studies. Then we consider the most recents estimates of the linear bias parameter $b_{\rm LIN}$ in this redshift range available in the literature and compare them with our value of $\hat{b}$.In these comparisons all results were rescaled to the value $\sigma_8 = 0.9$ adopted in this paper whenever required.

6.1. Galaxy bias from counts in cells

In Figure 16 we plot the values of $\hat{b}$ and $\tilde{b}/\hat{b}$ obtained from our analysis as a function of redshift (filled and open red dots) and compare these values to those obtained by M05 (green triangles) and K11 (blue squares) from counts in cells following a procedure similar to ours. We do not consider the results of the analyses of Simon et al. (2007) and Jullo et al. (2012) since these authors estimate the so-called correlation parameter that accounts for both non-linearity and stochasticity.

We only considered objects that, at a given redshift, span a similar range of magnitudes to avoid mixing evolution and luminosity dependence. For VIPERS we consider objects with $M_B < -19.1 - z - 5\log(h)$. For zCOSMOS we consider objects above a similar cut-off, $M_B = -19.22 - z - 5\log(h)$. For the VVDS-Deep sample, M05 use a redshift-independent luminosity threshold of $M_B = -20.0 - 5\log(h)$, which is comparable with the above cut-offs in the range $z = [0.8, 1.1]$. We considered an additional VIPERS subsample cut at the same constant magnitude limit as M05 to improve the consistency in the comparison with VVDS.

In the case of M05, the values of $\hat{b}$ and $\tilde{b}/\hat{b}$ shown in Figure 16 were inferred from the published values of $\tilde{b}$ and $\hat{b}/\tilde{b}$. In addition, M05 do not provide the errors for $\tilde{b}/\hat{b}$. The error bars shown in the plot were extrapolated from the errors on $\hat{b}$ under the assumption that the ratio of the errors on $\hat{b}$ and those on $\tilde{b}/\hat{b}$ are the same for the two datasets. The comparison between zCOSMOS and VIPERS shows that this assumption is approximately valid. The zCOSMOS points are plotted at the centre of their redshift bins. In the VVDS case we added an offset $\Delta_z = +0.02$ to avoid overlapping. Finally, we restrict our comparison to counts in cells of $R = 8\ h^{-1}$ Mpc since this is the minimum cell size considered by K11 and the only one common to the three analyses.

The values $\hat{b}$ of zCOSMOS galaxies (bottom panel) are in agreement with those of VIPERS galaxies. These values increase with redshift in both cases. This trend is more evident in the zCOSMOS case, while for VIPERS the evolution is detected only with a significance of $\sim 1\ \sigma$ only. Our results do not match those of M05 at $z = 0.8$, where the two samples overlap. The significance of the discrepancy, however, is about 1 σ. A similar mismatch was observed between VVDS-Deep zCOSMOS and interpreted by K11 in terms of different clustering amplitude in the two datasets (McCracken et al. 2007; Meneux et al. 2009; Kovač et al. 2011). Indeed, zCOSMOS is characterised by prominent structures and large spatial coherence as opposed to the VVDS Deep field. This difference was interpreted as a manifestation of cosmic variance. The VIPERS survey was designed to reduce the impact of cosmic variance and solve these types of controversies. In this specific case, the agreement between VIPERS and zCOSMOS galaxies suggests that the bias of the latter is closer to the cosmic mean than that of the VVDS-Deep field.

The comparison among the non-linear bias parameters of the three galaxy samples (upper panel of Fig. 16) corroborates this conclusion. The values of $\tilde{b}/\hat{b}$ for zCOSMOS and VIPERS galaxies agree with each other and significantly deviates from unity. Thanks to the smaller error bars in VIPERS these deviations are now detected with higher statistical significance. Deviations from non-linear bias in the VVDS-Deep are larger than in VIPERS but the statistical significance for this mismatch is just about 1 σ.

Figure 17 is analogous to Figure 16. It shows the values of $\hat{b}$ and $\tilde{b}/\hat{b}$ for galaxy subsamples extracted from VIPERS (red dots) and zCOSMOS (blue squares) using magnitude cuts brighter than before: $M_B = -19.7 - z - 5\log(h)$ for VIPERS and $M_B = -19.72 - z - 5\log(h)$ for zCOSMOS. Our results confirm those obtained with the fainter samples; the values of $\hat{b}$ and $\tilde{b}/\hat{b}$ for VIPERS galaxies agree with those of zCOSMOS galaxies in the redshift range in which the two analyses overlap. Non-linearity is detected at more than 1-σ in the VIPERS sample alone. No comparison was made with the VVDS-Deep sample in this case since none of the subsamples analysed by M05 match these luminosity cuts.

6.2. Linear bias from galaxy clustering

In Section 5 we saw that the bias of VIPERS galaxies deviates from linearity at all redshifts and on all scales explored. The amount of non-linearity, quantified by the parameter $\tilde{b}/\hat{b}$, is rather small, of the order of a few per cent. This means that $\hat{b}$ is reasonably similar to $b_{\rm LIN}$ and, therefore, can be compared with the linear bias parameter computed in other analyses.

In the following, we therefore compare the values of $\hat{b}$ computed in this work with the values of $b_{\rm LIN}$ obtained from different datasets in the same redshift range but using a variety of bias estimators. Galaxy bias at these redshifts has been estimated from both galaxy clustering and weak lensing. The latter probe, however, has either focused on bright objects used to trace baryonic acoustic oscillations (Comparat et al. 2013) or to explore bias dependence on the stellar mass (Jullo et al. 2012). Therefore, we

Table 2. Bias parameters of VIPERS galaxies and their errors.

z-range	M_B- cut $M_B(z=0) - 5\log h$	R h^{-1} Mpc	δ_{MAX}	$\hat{b}$	$\sigma^{\hat{b}}_{RND}$	$\sigma^{\hat{b}}_{CV}$	$\sigma^{\hat{b}}_{SYS}$	$\tilde{b}/\hat{b}$	σ^{NL}_{RND}	σ^{NL}_{CV}	σ^{NL}_{SYS}
0.5-0.7	$-18.6 - z$	4	11(15)	1.01 (0.98)	0.04	0.02	0.02	1.018 (1.021)	0.005	0.003	-0.003
0.5-0.7	$-19.1 - z$	4	11(15)	1.06 (1.03)	0.04	0.02	0.02	1.017 (1.021)	0.006	0.002	-0.004
0.5-0.7	$-19.5 - z$	4	11(15)	1.10 (1.07)	0.04	0.02	0.02	1.017 (1.020)	0.010	0.002	-0.005
0.5-0.7	$-19.9 - z$	4	11(15)	1.23 (1.20)	0.05	0.02	0.02	1.007 (1.009)	0.012	0.002	-0.006
0.7-0.9	$-19.1 - z$	4	11(15)	1.15 (1.12)	0.03	0.02	0.01	1.012 (1.018)	0.004	0.002	-0.002
0.7-0.9	$-19.5 - z$	4	11(15)	1.17 (1.14)	0.03	0.02	0.02	1.015 (1.019)	0.004	0.001	-0.002
0.7-0.9	$-19.9 - z$	4	11(15)	1.20 (1.15)	0.04	0.02	0.02	1.013 (1.018)	0.005	0.001	-0.002
0.9-1.1	$-19.5 - z$	4	11(15)	1.45 (1.45)	0.05	0.02	0.02	1.008 (1.007)	0.006	0.001	−0.002
0.9-1.1	$-19.9 - z$	4	11(15)	1.49 (1.46)	0.11	0.02	0.03	1.008 (1.008)	0.018	0.001	−0.003
0.5-0.7	$-18.6 - z$	6	10(14)	1.16 (1.14)	0.04	0.04	0.03	1.012 (1.014)	0.004	0.002	-0.003
0.5-0.7	$-19.1 - z$	6	10(14)	1.22 (1.21)	0.05	0.04	0.03	1.012 (1.014)	0.004	0.002	-0.003
0.5-0.7	$-19.5 - z$	6	10(14)	1.27 (1.26)	0.06	0.04	0.04	1.010 (1.011)	0.006	0.002	-0.004
0.5-0.7	$-19.9 - z$	6	10(14)	1.35 (1.34)	0.07	0.05	0.05	1.009 (1.011)	0.007	0.002	-0.004
0.7-0.9	$-19.1 - z$	6	10(14)	1.35 (1.34)	0.03	0.03	0.02	1.005 (1.005)	0.003	0.002	-0.002
0.7-0.9	$-19.5 - z$	6	10(14)	1.37 (1.36)	0.03	0.02	0.02	1.007 (1.008)	0.003	0.002	-0.002
0.7-0.9	$-19.9 - z$	6	10(14)	1.43 (1.41)	0.03	0.03	0.02	1.009 (1.011)	0.004	0.002	-0.002
0.9-1.1	$-19.5 - z$	6	10(14)	1.85 (1.84)	0.05	0.03	0.01	1.007 (1.008)	0.004	0.001	-0.002
0.9-1.1	$-19.9 - z$	6	10(14)	1.85 (1.84)	0.06	0.03	0.02	1.011 (1.012)	0.004	0.001	-0.002
0.5-0.7	$-18.6 - z$	8	8(12)	1.24 (1.23)	0.07	0.05	0.04	1.007 (1.008)	0.005	0.003	-0.003
0.5-0.7	$-19.1 - z$	8	8(12)	1.32 (1.31)	0.07	0.05	0.04	1.007 (1.008)	0.004	0.002	-0.003
0.5-0.7	$-19.5 - z$	8	8(12)	1.36 (1.36)	0.08	0.06	0.05	1.007 (1.008)	0.005	0.002	-0.003
0.5-0.7	$-19.7 - z$	8	8(12)	1.40 (1.39)	0.09	0.06	0.06	1.006 (1.007)	0.005	0.002	-0.003
0.5-0.7	$-19.9 - z$	8	8(12)	1.44 (1.44)	0.09	0.06	0.06	1.006 (1.006)	0.005	0.002	-0.003
0.7-0.9	$-19.1 - z$	8	8(12)	1.44 (1.44)	0.04	0.04	0.02	1.003 (1.003)	0.003	0.002	-0.002
0.7-0.9	$-19.5 - z$	8	8(12)	1.46 (1.46)	0.05	0.04	0.03	1.005 (1.005)	0.003	0.002	-0.002
0.7-0.9	$-19.7 - z$	8	8(12)	1.48 (1.48)	0.07	0.04	0.03	1.008 (1.008)	0.004	0.002	-0.002
0.7-0.9	$-19.9 - z$	8	8(12)	1.51 (1.50)	0.08	0.04	0.04	1.008 (1.009)	0.006	0.002	-0.002
0.9-1.1	$-19.5 - z$	8	8(12)	2.01 (2.01)	0.06	0.04	0.02	1.010 (1.010)	0.003	0.002	-0.002
0.9-1.1	$-19.7 - z$	8	8(12)	1.98 (1.98)	0.06	0.04	0.03	1.009 (1.009)	0.003	0.003	-0.002
0.9-1.1	$-19.9 - z$	8	8(12)	2.01 (2.01)	0.07	0.05	0.04	1.015 (1.015)	0.004	0.003	-0.001
0.7-0.9	−20.0	8	8(12)	1.43 (1.42)	0.04	0.03	0.02	1.006 (1.005)	0.002	0.002	-0.001

Notes. Col. 1: redshift range. Col. 2: z-dependent B-band magnitude cut. Col. 3: cell radius [h^{-1} Mpc]. Col. 4: maximum over-density considered in the analysis δ_{MAX}; the value in parenthesis indicates $\bar{\delta}_{MAX}$. Col. 5: estimated value of the bias moment $\hat{b}$; the values in parenthesis refer to measurements performed at $\bar{\delta}_{MAX}$. Col. 6: total random error on $\hat{b}$. Col. 7: cosmic variance contribution to $\hat{b}$ error. Col. 8: systematic error on $\hat{b}$. Col. 9: estimated value of the non-linearity parameter $\tilde{b}/\hat{b}$; values in parenthesis refer to measurements performed at $\bar{\delta}_{MAX}$. Col. 10: total random error on $\tilde{b}/\hat{b}$. Col. 11: cosmic variance contribution to $\tilde{b}/\hat{b}$ error. Col. 12: systematic error on $\tilde{b}/\hat{b}$.

focus on the values of b_{LIN} obtained from galaxy clustering in other datasets available in the literature.

The results of these comparisons are shown in Figure 18 in which we plot the most recent estimates of both $\hat{b}$ and b_{LIN} in the three redshift bins as a function of the magnitude cut. We consider the reference scale of $R = 8\ h^{-1}$ Mpc since this is the size of the cells used to measure $\hat{b}$ in VIPERS (large red circles), VVDS-Deep (small orange pentagons), and zCOSMOS (small, light green circles), as shown in the previous section. Magnitudes on the X-axis are specified in the B-band since this band is used in most of the considered samples with the exception of the PRIMUS and CHFTLS-wide. For these two latter cases, we consider the g-band magnitude and transform it into B-band according to the $g-$B versus z relation measured from the VIPERS catalogue. Finally, all results were normalised to $\sigma_8 = 0.9$.

The large red circles represent the $\hat{b}$ values obtained from VIPERS, for the systematic errors listed in Table 2. Therefore these values are slightly different from those shown in Figures 16 and 17.

The light blue asterisks represent the b_{LIN} values obtained from the Wide part of the Canada-France-Hawaii Legacy Survey (CFHTLS) (Coupon et al. 2012). In this case, the bias values were computed from $\sim 3 \times 10^6$ galaxies in the redshift interval $z = [0.2, 1.2]$ by fitting a Halo Occupation Distribution model to the measured angular correlation function. These b_{LIN} values were obtained by integrating the halo bias over the halo mass function. Therefore they are integral quantities much like $\hat{b}$, which is computed by integrating over the mass PDF (Eq. 2). The b_{LIN} values of CFHTLS galaxies agree well with our results in all redshift bins, including $z = [0.9, 1.1]$.

The brown crosses in the middle panel of Figure 18 are from Skibba et al. (2014) and show the bias of galaxies in the PRIMUS catalogue. The PRIMUS (Coil et al. 2011) galaxy survey is carried out using a low-resolution spectrograph and complete down to $i < 23$. This dataset, which covers five independent fields (including the COSMOS field), spans the redshift range $z = [0.2, 1.0]$. Here we focus on the interval $z = [0.5, 1.0]$ and plot the corresponding bias values in the middle panel. The bias was estimated from the projected galaxy 2-point correlation function, $w_{p,g}(r)$, as $b_{LIN}(r) = \sqrt{w_{p,g}(r)/w_{p,m}(r)}$, where the projected 2-point correlation function of the matter, $w_{p,m}(r)$, was modelled assuming a flat ΛCDM model with $\sigma_8 = 0.8$. The

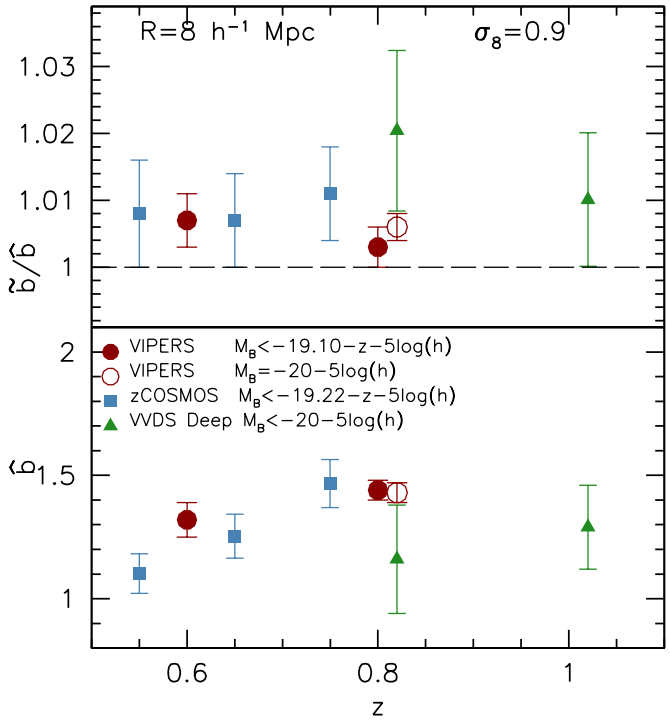

Fig. 16. Comparison among the values of $\tilde{b}/\hat{b}$ (top panel) and $\hat{b}$ (bottom panel) for zCOSMOS galaxies (blue squares) for VVDS-Deep galaxies (green triangles) and VIPERS galaxies (filled red circles). All samples are luminosity limited and the magnitude cuts are indicated in the plot. The open red circles represents a VIPERS subsample matching the magnitude cut and redshift range of the VVDS-Deep sample. Estimates for the bias parameters of zCOSMOS are taken from K11, and those for VVDS-Deep galaxies are from M05.

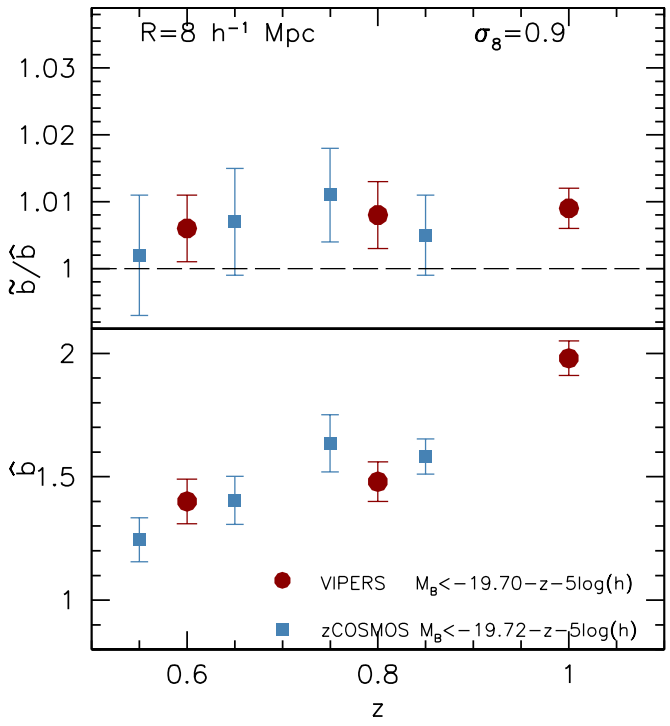

Fig. 17. Same as figure 16, but referring to brighter VIPERS ($M_B < -19.7 - z - 5\log(h)$, red dots) and zCOSMOS galaxies ($M_B < -19.72 - z - 5\log(h)$, blue squares).

bias of PRIMUS galaxies is systematically larger than that of VIPERS. However, the significance of the mismatch is below *1-σ*.

The purple hexagons in the plot show $b_{\rm LIN}$ of ALHAMBRA galaxies (Arnalte-Mur et al. 2014). The photometric redshift survey ALHAMBRA covers seven independent fields, including DEEP2 and COSMOS. Photometric redshifts are accurate enough to measure the projected galaxy correlation function at different redshifts and, from this, to estimate the bias. In Figure 18 we show the $b_{\rm LIN}$ values estimated in three redshift bins: $[0.35, 0.65]$ (top panel), $[0.55, 0.85]$ (middle panel), and $[0.75, 1.05]$ (bottom panel). We did not consider the interval $z = [0.95, 1.25]$ since it is largely beyond the VIPERS range. We show two sets of points. Small open hexagons represent the values of $b_{\rm LIN}$ obtained from the clustering of galaxies in all seven fields (labelled ALHAMBRA+ in the plot). Filled hexagons (labelled ALHAMBRA-) illustrate the effect of removing two "outlier" fields, COSMOS and ELAIS-N1, which are characterised by a high degree of clustering. The bias of galaxies in ALHAMBRA- agrees with that of VIPERS for $z < 0.9$. In the last redshift bin, for $M_B(z = 1) < -20.56 - 5\log(h)$ the bias of ALHAMBRA- is ~ 1.5-σ below that of VIPERS. However, the discrepancy disappears when one considers ALHAMBRA+ and seems to reappear, with a reverse sign, at higher luminosities.

The green triangles show the $b_{\rm LIN}$ values obtained from the projected galaxy 2-point correlation function of galaxies brighter than $M_B - 5\log(h) = -20.5 - 5\log(h)$ at $z = [0.9, 1.1]$ in the DEEP2 survey (Coil et al. 2006). In the brightest magnitude bin, where the three samples overlap, we find that the bias of DEEP2 galaxies is significantly smaller than that of VIPERS and ALHAMBRA objects.

To summarise, we find a good agreement between the value of $\hat{b}$ measured in our work and those of $b_{\rm LIN}$ estimated in a number of surveys in the range $z = [0.5, 0.9]$. In particular, our measurements agree with those of K11 (small, light green circles) who used the same technique to estimate $\hat{b}$.

In the outermost redshift shell not all the bias values measured in different surveys agree with each other. The value of $b_{\rm LIN}$ for DEEP2 and, to a lesser extent, for ALHAMBRA- galaxies, are smaller than $\hat{b}$ from VIPERS. This mismatch may indicate either a genuine difference in the clustering properties of the different samples or deviations from linear bias highlighted by the different bias estimators.

To quantify the impact of non-linear bias, we compare our $\hat{b}$ values with the corresponding $b_{\rm LIN}$ estimated from the very same VIPERS subsamples considered here. Figure 19 compares $\hat{b}$ from VIPERS (red filled symbols) with $b_{\rm LIN}$ from Marulli et al. (2013) (blue filled squares, also shown, for reference in Figure 18). The two estimates agree at all redshifts but the last redshift bin where the bias of Marulli et al. (2013) matches that of DEEP2 galaxies and, consequently, is significantly below our $\hat{b}$ value.

Like most of the other measurements, Marulli et al. (2013) estimated $b_{\rm LIN}$ from the projected 2-point correlation function. More precisely, they averaged the correlation signal over the interval $r = [1, 10]\, h^{-1}$ Mpc . In the presence of a scaled dependent bias, a manifestation of which is a $\tilde{b}/\hat{b}$ ratio different from unity, it is not obvious which effective scale of the bias is estimated by Marulli et al. (2013). In our comparison we implicitly assumed that this scale is the same as the cell size, i.e. $8\, h^{-1}$ Mpc . In fact, a small scale seems more appropriate, especially if one accounts for the fact that errors in the projected correlation function increases with the pair separation. For this reason, we also show $\hat{b}$

measured in cells of $R = 6\ h^{-1}$ Mpc (orange hexagons). In this case, the significance of the mismatch is significantly reduced. Decreasing the scale to $R = 4\ h^{-1}$ Mpc (not shown) would bring the two values into agreement at the price, however, of creating a mismatch at lower redshifts.

Focusing on the VIPERS sample, a more homogeneous comparison can be performed considering the $b_{\rm LIN}$ value obtained by Cappi et al. (2015) from counts in cells of $R = 8\ h^{-1}$ Mpc (brown asterisks, in the plots). In this case the results agree with ours within the (rather large) error bars.

In the figure we also show the VIPERS linear bias estimated by Granett et al. (2015) (green triangles) from a Bayesian reconstruction of a Wiener filtered, adaptively smoothed galaxy density field. The result agrees with that of Marulli et al. (2013). However, as in that case, it is difficult to associate an effective scale to the filtering procedure and perform a homogeneous comparison with our estimate.

Therefore, all the bias estimates of the VIPERS galaxies agree with each other at $z < 0.9$, a sign that galaxy bias is largely independent of scales. At higher redshifts we observe some discrepancies among the various estimates whose significance, however, is difficult to assess since the different estimates are sensitive to different scales. It is safe to conclude that the scale-dependence bias of VIPERS galaxies is more pronounced at high redshifts, as confirmed by the results presented in Section 5.3, and that this can account for most of the discrepancies seen in Fig. 19.

An additional, though minor, source of discrepancy is incompleteness. At $z \sim 1$ the 90 % completeness limit in VIPERS is $M_B - 5\log(h) \sim -21.0$ for red galaxies and about half a magnitude fainter for the blue galaxies. Since red galaxies are more biased than the blue galaxies, selecting objects at this luminosity cut underestimates the bias of the composite VIPERS sample. The amplitude of the effect depends on the luminosity cut. We conclude that deviations from linear bias cannot be ignored at high redshifts and that using $b_{\rm LIN}$ as a proxy for galaxy bias leads to significant systematic errors.

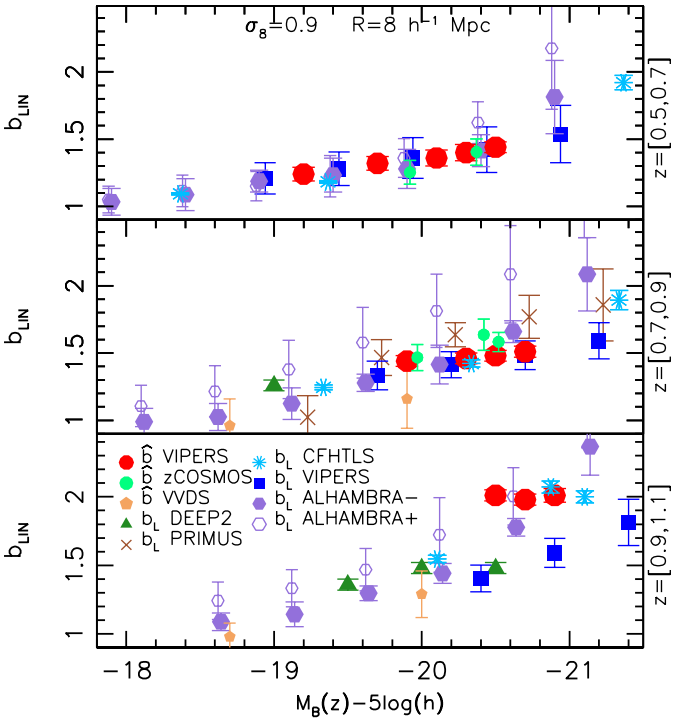

Fig. 18. Comparison between the bias parameters $\hat{b}$ and b_{LIN} obtained from galaxy counts and galaxy clustering, respectively. Large red circles: $\hat{b}$ of VIPERS galaxies. Blue squares: b_{LIN} of VIPERS from Marulli et al. (2013). Green triangles: b_{LIN} for DEEP2 galaxies from Coil et al. (2006). Brown crosses: b_{LIN} for PRIMUS galaxies from Skibba et al. (2014). Light blue asterisks: b_{LIN} for CHFTLS galaxies from Coupon et al. (2012). Purple hexagons: b_{LIN} for ALHAMBRA galaxies from Arnalte-Mur et al. (2014).Small light green dots: $\hat{b}$ for zCOSMOS galaxies from Kovač et al. (2011). Small light brown pentagons: $\hat{b}$ for VVDS-Deep galaxies from Marinoni et al. (2005). Values of $\hat{b}$ were measured on a scale $R = 8\ h^{-1}$ Mpc .

7. Discussion and conclusions

The importance of characterising galaxy bias at intermediate redshifts stems from the need to infer the properties of the distribution of the mass from that of the galaxies. This will be especially important in future redshift surveys aimed at an accurate estimate of the cosmological parameters. This has prompted several efforts to estimate galaxy bias at $z > 0.5$ exploiting weak lensing, galaxy clustering, and galaxy counts. Most of these works assume that galaxy bias is linear and deterministic and provide an estimate for the linear bias parameter. In this work we questioned this assumption and searched for possible deviations from linear bias. This issue has already been investigated by M05 and K11 (using counts in cells and significantly smaller samples) and by Simon et al. (2007) and Jullo et al. (2012) with conflicting evidence, as discussed in the introduction.

Our work builds upon these results improving the original strategy of M05 and K11 in several aspects. First of all, it is based on a new dataset of $\sim 50,000$ galaxies distributed over a much larger volume than its predecessors. This significantly reduces the impact of cosmic variance that in previous studies dominated the error budget. Secondly, we use a new technique to infer the mean biasing function from counts in cells that, under the hypothesis of local Poisson sampling, accounts and automatically corrects for shot noise. This improvement greatly increases our ability to recover the biasing function since Poisson noise is the main source of stochasticity in the bias relation. Thirdly, owing to the size of the sample, we are able to explore the bias dependence on magnitude, redshift, and scale with unprecedented accuracy. We postpone the investigation of additional dependences on galaxy colour and stellar mass to a future analysis to be performed with the final VIPERS sample and new mock galaxy catalogues designed to mimic these galaxy properties.

The main results of our study are:

The overall qualitative behaviour of the mean biasing function of VIPERS galaxies is similar to that of zCOSMOS and VVDS-Deep galaxies as well as to that of the synthetic VIPERS galaxies in the mock catalogues that we used to estimate errors. The shape of the mean biasing function is close to linear in regions above the mean density. It deviates from linear bias at $\delta < 0$. More specifically, above the threshold $\delta_{\rm TH}$ at which $\delta_g = -0.9$, the bias function is significantly steeper than its mean slope $\hat{b}$ on scales smaller than $8\ h^{-1}$ Mpc . For over-densities below $\delta_{\rm TH}$ the mean biasing function features a tail that cannot be accounted for by linear biasing. The over-density threshold $\delta_{\rm TH}$ has been interpreted as a typical density scale below which very few galaxies form. In our analysis, we find that this threshold increases with the redshift and with the luminosity cut-of,f so that at moderate redshifts low-density regions are typically populated by faint galaxies.

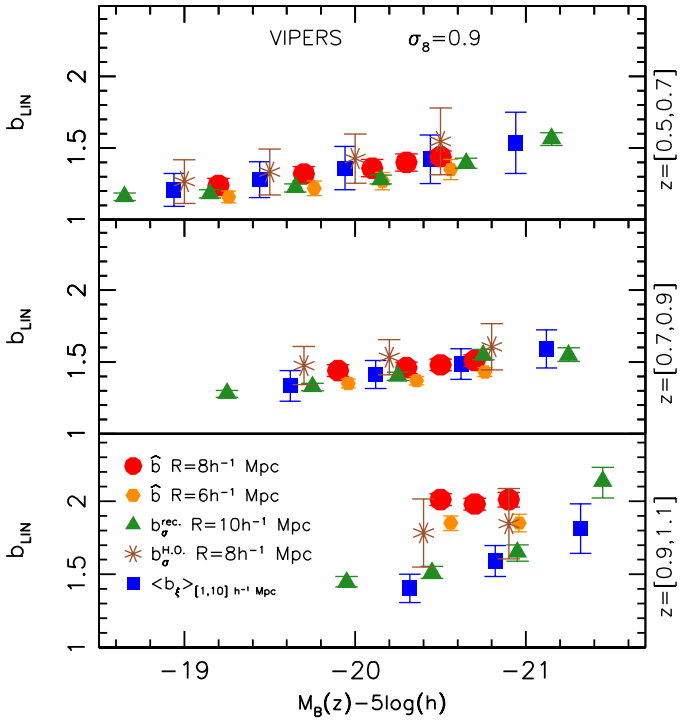

Fig. 19. Comparison between the bias parameter $\hat{b}$ obtained from our analysis on a scale of $R = 8\ h^{-1}$ Mpc (large red circles), $R = 6\ h^{-1}$ Mpc (small orange hexagons) and the linear bias parameters of VIPERS galaxies, b_{LIN}, obtained by Marulli et al. (2013) (blue squares), Cappi et al. (2015) (brown asterisks), and Granett et al. (2015) (green triangles).

The biasing function shows small but significant deviations from linearity at all redshifts, scales, and magnitude intervals that we explored. The parameter $\tilde{b}/\hat{b}$ that we use to quantify non-linearity neither seems to evolve with the redshift nor to depend on the luminosity. A scale dependence is observed at low redshifts below $6\ h^{-1}$ Mpc with only $\sim 1\,\sigma$ significance.

We confirm that galaxy bias depends on luminosity. The mean slope $\hat{b}$ of the biasing function, a good proxy to linear bias given the small degree of non-linearity, increases with the luminosity threshold used to select the galaxy sample. The effect is significant for $z = [0.5, 0.7]$, probably thanks to the large magnitude leverage here compared to the bin $z = [0.9, 1.1]$, in which we can probe a much smaller magnitude range $\Delta M_B = 0.4$. The value of δ_{TH} also increases with the magnitude, suggesting that the efficiency of galaxy formation decreases with the luminosity of the object.

We confirm that galaxy bias increases with redshifts as predicted by most bias models and verified in other datasets. In our case we find evidence for a rapid evolution beyond $z = 0.9$. This result is highly significant and robust since it depends neither on the scale nor on the luminosity of the objects. The statistical significance of this result depends on the reliability of our error analysis, which is based on a mock galaxy catalogue designed to match the correlation properties of VIPERS galaxies. We verified that mock catalogues are very realistic in the sense that their biasing function matches that of real objects remarkably well.In this analysis, we modelled all known sources of systematic errors, including the magnitude dependence of the spectroscopic sampling rate that was not originally included in the mock catalogues. We find no evidence for systematic errors that might mimic a spurious evolution in the bias moments $\hat{b}$ and $\tilde{b}$.

The value of $\hat{b}$ increases with the scale from 4 to 8 h^{-1} Mpc .We interpret this in the framework of the halo model as the transition between the one-halo and two-halo contribution to galaxy bias. The same trend is seen in our mock catalogues, in which objects were extracted assuming the HOD model and in previous analyses performed at lower redshifts using both galaxy clustering (e.g. Zehavi et al. (2005)) and weak lensing (e.g. Hoekstra et al. (2002), Simon et al. (2007)). This is the first time that this effect is detected at high redshifts with counts in cells statistics.

We compared our results with those of M05 and K11. These authors performed an analysis similar to that presented here in a similar range of redshifts. We limited the comparison on a scale of $8\ h^{-1}$ Mpc , which is common to the three analyses. We find that the values of $\hat{b}$ of VIPERS and zCOSMOS galaxies agree within the errors. M05 find a smaller degree of biasing but the difference is of the order of $1-\sigma$. We conclude that the claimed discrepancy between K11 and M05 results is a manifestation of cosmic variance.

Deviations from linear biasing were also detected by M05 and K11, although with a lower significance than in our case. Our results agree with those of K11. In M05 the degree of non-linearity is slightly larger than in our case but the discrepancy is barely larger than 1-σ. The bias non-linearity is sometimes expressed in terms of the parameter b_2 of the second-order Tailor expansion of δ (Fry & Gaztanaga 1993). Cappi et al. (2015) analysed the same VIPERS dataset using higher order statistics, which is a procedure that is less sensitive to non-linear bias than ours. They detected a deviation from linear bias at $z \leq 0.9$ with a significance of $\sim$ 1-σ. Their b_2 value turned out to be negative, in agreement with M05 and Marinoni et al. (2008) and, qualitatively, with our results too.

measured from the clustering of galaxies in recent galaxy redshift surveys (DEEP2, PRIMUS, CHFTLS-wide, and ALHAMBRA). This comparison is qualitative since it assumes that bias is linear, while our analysis has detected a small, but significant, degree of non-linearity in the bias of VIPERS galaxies. The comparison is generally successful at $z < 0.9$, where we find a very good agreement with all existing results. In this redshift range our results provide additional evidence in favour of a luminosity-dependent bias and of a weak evolution. At $z > 0.9$, where the spread among current results is large, our results favour the case of a significant bias evolution, in agreement with the CHFTLS-wide and ALHAMBRA analyses.

Our results confirm the importance of going beyond the simplistic linear biasing hypothesis. Galaxy bias is a complicated phenomenon. It can be non-deterministic, non-local, and non-linear. In this work we focused on deviations from linearity under the assumption that stochasticity is dominated by (and, consequently, accounted for) Poisson noise and that non-local effects are smoothed out within the volume of our cells. While the validity and the impact of these assumptions can (and will) need to be tested, our results show that the application of an improved statistical tool to the new VIPERS dataset is already able to detect deviations from linear bias with 5 – 10 % accuracy.

Acknowledgments

We acknowledge the crucial contribution of the ESO staff for the management of service observations. In particular, we are deeply grateful to M. Hilker for his constant help and sup-

port of this programme. Italian participation in VIPERS is funded by INAF through PRIN 2008 and 2010 programmes. LG and BRG acknowledge support of the European Research Council through the Darklight ERC Advanced Research Grant (# 291521). OLF acknowledges support of the European Research Council through the EARLY ERC Advanced Research Grant (# 268107). Polish participants are supported by the Polish Ministry of Science (grant N N203 51 29 38), the Polish-Swiss Astro Project (co-financed by a grant from Switzerland, through the Swiss Contribution to the enlarged European Union), the European Associated Laboratory Astrophysics Poland-France HECOLS, and a Japan Society for the Promotion of Science (JSPS) Postdoctoral Fellowship for Foreign Researchers (P11802). GDL acknowledges financial support from the European Research Council under the European Community's Seventh Framework Programme (FP7/2007-2013)/ERC grant agreement n. 202781. WJP and RT acknowledge financial support from the European Research Council under the European Community's Seventh Framework Programme (FP7/2007-2013)/ERC grant agreement n. 202686. WJP is also grateful for support from the UK Science and Technology Facilities Council through the grant ST/I001204/1. EB, FM, and LM acknowledge the support from grants ASI-INAF I/023/12/0 and PRIN MIUR 2010-2011. LM also acknowledges financial support from PRIN INAF 2012. YM acknowledges support from CNRS/INSU (Institut National des Sciences de l'Univers) and the Programme National Galaxies et Cosmologie (PNCG). CM is grateful for support from specific project funding of the Institut Universitaire de France and the LABEX OCEVU. MV is supported by FP7-ERC grant "cosmoIGM" and by PD-INFN INDARK. CDP wishes to thank Pangea Formazione S.r.l. for supporting her in the final stages of this work.

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