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Erratum

The X-ray emission from Young Stellar Objects in the ρ Ophiuchi cloud core as seen by XMM-Newton

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Key words. open clusters and associations: individual: ρ Ophiuchi cloud – stars: pre-main sequence – stars: low-mass, brown dwarfs – X-rays: stars – infrared: stars – errata, addenda

We wrote in the Sect. 5 of Ozawa et al. (2005) that with our *XMM-Newton* observation, "we detect (the bona fide brown dwarf) GY 141 (ROXN-10) with a count rate of 7.7 cts ks^{-1}, i.e., at a level 90 times higher than during the (previous) *Chandra* observation (where only 8 X-ray photons were collected during the 100 ks exposure)". Indeed this factor of 90 is the ratio of the count rates observed by *XMM-Newton* and *Chandra*. To obtain the *absolute* increase of brightness of GY 141, it has to be divided by 6.8 – the median of the ratios between *XMM-Newton* and *Chandra* count rates for X-ray sources detected by both observatories (see Sect. 3 of Ozawa et al. 2005) – which leads to ∼14. Consequently, the increase of brightness of GY 141 that should appear in the abstract and the summary of Ozawa et al. (2005), is *one* order of magnitude. This doesn't change the conclusion of our article.

References

Ozawa, H., Grosso, N., & Montmerle, T. 2005, A&A, 429, 963