

## Transiting exoplanets from the CoRoT space mission<sup>★</sup>

### XX. CoRoT-20b: A very high density, high eccentricity transiting giant planet

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#### ABSTRACT

We report the discovery by the CoRoT space mission of a new giant planet, CoRoT-20b. The planet has a mass of  $4.24 \pm 0.23 M_{\text{Jup}}$  and a radius of  $0.84 \pm 0.04 R_{\text{Jup}}$ . With a mean density of  $8.87 \pm 1.10 \text{ g cm}^{-3}$ , it is among the most compact planets known so far. Evolutionary models for the planet suggest a mass of heavy elements of the order of  $800 M_{\oplus}$  if embedded in a central core, requiring a revision either of the planet formation models or both planet evolution and structure models. We note however that smaller amounts of heavy elements are expected by more realistic models in which they are mixed throughout the envelope. The planet orbits a G-type star with an orbital period of 9.24 days and an eccentricity of 0.56. The star's projected rotational velocity is  $v \sin i = 4.5 \pm 1.0 \text{ km s}^{-1}$ , corresponding to a spin period of  $11.5 \pm 3.1$  days if its axis of rotation is perpendicular to the orbital plane. In the framework of Darwinian theories and neglecting stellar magnetic breaking, we calculate the tidal evolution of the system and show that CoRoT-20b is presently one of the very few Darwin-stable planets that is evolving toward a triple synchronous state with equality of the orbital, planetary and stellar spin periods.

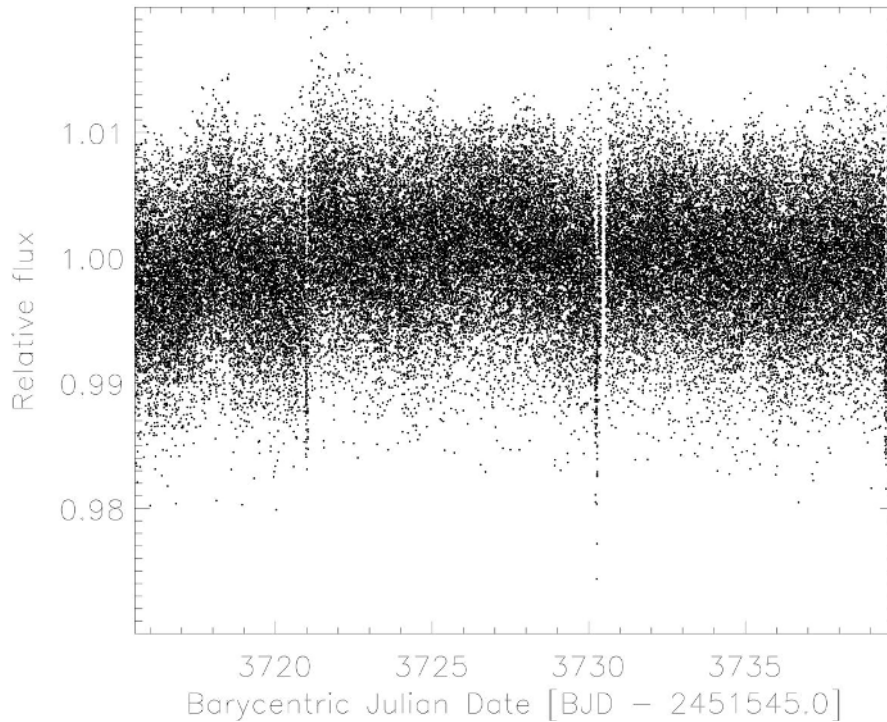
**Key words.** planetary systems – stars: fundamental parameters – techniques: photometric – techniques: radial velocities – techniques: spectroscopic

### 1. Introduction

The existence of a planet population at very short orbital distance of a star, typically  $a < 0.1 \text{ AU}$ , with its wide range of

orbital and physical properties is an intriguing phenomenon. In-situ formation of these massive bodies so close to their host star at a location where the circumstellar material is depleted and warm, indeed appears highly unlikely. Planet migration from further away distances where solid material is abundant, triggered by gravitational interactions, is invoked to account for this population. The exact process responsible remains unclear but two major mechanisms have been proposed: gradual planet

<sup>★</sup> The CoRoT space mission, launched on December 27th 2006, has been developed and is operated by CNES, with the contribution of Austria, Belgium, Brazil, ESA (RSSD and Science Programme), Germany, and Spain.



**Fig. 1.** The 24.278-day long CoRoT-20 reduced light curve at a constant 512-s time sampling.

migration caused by interaction with the circumstellar gas disk (Lin et al. 1996; Papaloizou et al. 2007) or planet-planet or planet-companion star interactions combined with tidal dissipation (e.g., Rasio & Ford 1996; Fabrycky & Tremaine 2007; Nagasawa et al. 2008). Regardless of the exact nature of the formation path, these planets have undergone significant orbital evolution since the time of their formation. Their current properties provide in valuable constraints to help our understanding of their orbital evolution, especially the planet’s orbital eccentricity and the spin-orbit alignment of the system.

Until now, transit surveys have been more sensitive to planets of a very short orbital period. The situation improved thanks to the extended temporal coverage of ground-based transit surveys and the advent of space missions, CoRoT (Baglin et al. 2009; Deleuil et al. 2011) and *Kepler* (Borucki et al. 2010). As a consequence, the number of planets with orbital periods longer than a few days has significantly increased over the past two years. While the mean eccentricity of the close-in planets is close to zero, the transiting giant planets at greater orbital distance have a much wider range of eccentricity, a picture more consistent with the sample of planets found by radial velocity surveys. These trends imply that there is a third-body-induced migration with tidal circularization of an initial eccentric and possibly high-obliquity orbit (e.g., Winn et al. 2010; Matsumura et al. 2010; Pont et al. 2011). A consequence of this orbital evolution is the tidal destruction of the planet, which spirals down onto the star in the lifetime of the system (Gonzalez 1997; Jackson et al. 2009), a dramatic destruction that appears to be the fate of the vast majority of the transiting planets (Matsumura et al. 2010).

In this paper, we report the discovery of CoRoT-20b, a new member of the hot-Jupiter class population. The planet transits its G-type parent star every 9.24 days, along an orbit with a high eccentricity. The CoRoT observations are presented in Sect. 2. The accompanying follow-up observations and their results are described in Sects. 3 and 5 for the host star analysis. The final system parameters are derived in Sect. 4. We then discuss the

properties of this new planet in Sect. 6. We investigate its orbital evolution and fate but also its internal structure that raises new questions about the nature of such a compact object.

## 2. CoRoT-20b light curve

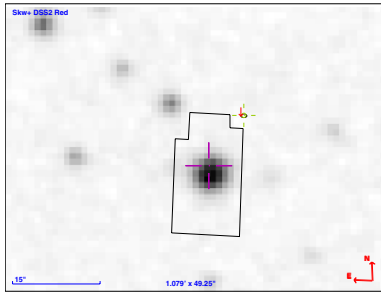
The planet was discovered in one of the fields observed by the CoRoT satellite (Baglin et al. 2009; Deleuil et al. 2011) in the so-called *anti-center* direction. This field, labeled SRa03, was monitored for 24.278 days starting on 1 March 2010. As a consequence of the DPU1 break down that took place in March 2009, the number of targets actually photometrically monitored by the instrument was reduced to a maximum of 6000 stars only. The released telemetry is used to oversample a much larger number of targets than initially possible, up to magnitude  $\approx 15$ . While not among the brightest stars of the field (Table 1), CoRoT-20 could be studied more effectively thanks to this new opportunity and its observation was performed with the regular 32 s time sampling. It was also bright enough to allow for three-color photometry. Two transits were detected in its light curve by the *Alarm mode* pipeline (Surace et al. 2008). The target was flagged as a good planetary candidate and placed among the highest priorities for follow-up observations.

The light curve of CoRoT-20b is displayed in Fig. 1. It shows a star rather quiet photometrically speaking, with no special prominent features. Three transits are clearly visible with a period of 9.24 days and a depth slightly shallower than 1%. For the detailed analysis, we used the light curve produced by the CoRoT calibration pipeline. It corrects for the main instrumental effects such as the CCD zero offset and gain, the background light, and the spacecraft jitter (see Auvergne et al. 2009). Portions of the light curve that were flagged by the pipeline as affected by particle impacts during the South Atlantic Anomaly crossing, were removed and ignored in the analysis. In total, the light curve consists of 56 869 photometric measurements and has a corresponding duty cycle of 88%.

**Table 1.** CoRoT-20b IDs, coordinates, and magnitudes.

CoRoT window ID	SRa03_E2_0999	
CoRoT ID	315239728	
USNO-B1 ID	0902-0091920	
2MASS ID	06305289+0013369	
GSC2.3 ID		
Coordinates		
RA (J2000)	97.720434	
Dec (J2000)	0.22692	
Magnitudes		
Filter	Mag	Error
$B^a$	15.31	
$V^a$	14.66	
$J^b$	12.991	0.023
$H^b$	12.652	0.026
$K^b$	12.512	0.027

**Notes.** <sup>(a)</sup> from USNO-B1 – Provided by Exo-Dat (Deleuil et al. 2009); <sup>(b)</sup> from 2-MASS catalog.



**Fig. 2.** Image of the DSS showing CoRoT-20 and its environment. The photometric mask used for CoRoT observations is overlotted on the target.

### 3. Follow-up observations

A photometric time-series of the star was obtained at the Wise observatory on November 14, 2010 to check whether an unknown nearby eclipsing binary could be the source of the transits (Deeg et al. 2009). The detection of a transit ingress excluded this configuration at the spatial resolution of Wise. The observed time of the ingress, with first contact at  $2\,455\,515.510 \pm 0.007$  HJD was then used to refine the period of CoRoT-20b, toward the value quoted in Table 3. Ground-based images from both Wise and the DSS show that the star is rather isolated (Fig. 2). This supports the very low contamination rate that was derived for the star within the CoRoT photometric mask (see Sect. 4).

Radial velocity (RV) observations started during the same season on December 9, 2010. We used the HARPS spectrograph (Mayor et al. 2003) mounted on the 3.6-m ESO telescope (Chile) as part of the ESO large program 184.C-0639, the SOPHIE spectrograph (Perruchot et al. 2008) on the 1.93-m telescope at the Observatoire de Haute Provence (France) and the FIES spectrograph on the Nordic Optical Telescope (Frandsen & Lindberg 1999) based on the 2.56-m Nordic Optical telescope in La Palma (Spain) under observing program P42-216.

We used the same instrument set-up as for previous CoRoT candidates follow-up: high resolution mode for HARPS and high efficiency mode for SOPHIE without acquisition of the simultaneous thorium-argon calibration, the second fiber being used to

**Table 2.** Log of radial velocity observations.

Date	HJD	$v_{\text{rad}}$ km s <sup>-1</sup>	$\sigma_{v_{\text{rad}}}$ km s <sup>-1</sup>	Spectrograph
2010-12-09	55 540.70645	60.728	0.0169	HARPS
2010-12-14	55 545.75533	60.353	0.0138	HARPS
2011-01-12	55 574.60927	60.453	0.0119	HARPS
2011-01-16	55 578.64264	60.985	0.0285	HARPS
2011-01-21	55 583.64035	60.443	0.0334	HARPS
2011-01-28	55 590.66159	60.320	0.0234	HARPS
2011-01-16	55 578.41934	60.813	0.0385	SOPHIE
2011-01-17	55 579.41057	61.048	0.0532	SOPHIE
2011-02-04	55 597.44200	60.948	0.0213	SOPHIE
2011-02-05	55 598.38017	61.077	0.0175	SOPHIE
2011-02-06	55 599.39134	60.267	0.0183	SOPHIE
2011-01-06	55 568.55091	60.611	0.032	FIES
2011-01-08	55 570.59973	61.090	0.060	FIES
2011-01-18	55 580.60213	60.260	0.036	FIES
2011-01-19	55 581.56506	60.146	0.034	FIES

monitor the Moon background light (Santerne et al. 2011). For HARPS and SOPHIE, the exposure time was set to one hour. We reduced data and computed RVs with the pipeline based on the weighted cross-correlation function (CCF) using a numerical G2 mask (Baranne et al. 1996; Pepe et al. 2002).

The FIES observations were performed in high-resolution mode with the 1.3 arcsec fiber yielding a resolving power  $R \approx 67\,000$  and a spectral coverage from 3600 to 7400 Å. Three consecutive exposures of 1200 s were obtained for each observation. Long-exposed ThAr spectra were acquired right before and after each science spectra set, as described in Buchhave et al. (2010). Standard IRAF routines were used to reduce, combine, and wavelength calibrate the nightly spectra. The RV measurements were derived by cross-correlating the science spectra with the spectrum of the RV standard star HD 50692 (Udry et al. 1999), observed with the same instrument set-up as CoRoT-20.

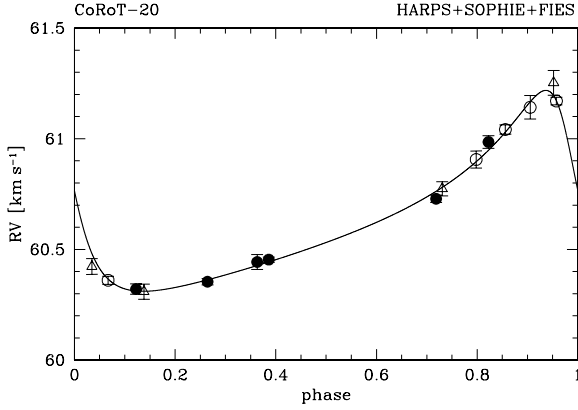
The 15 radial velocities of CoRoT-20b are listed in Table 2 and displayed in Fig. 3. They present a clear variation, in phase with the CoRoT ephemeris and consistent with a companion in the planet-mass regime with an eccentric orbit. Nevertheless, we investigated the possibility that an unresolved eclipsing binary was the source of observed transits. We performed the line-bisector analysis of the CCFs (see Fig. 4) and also verified that there is no dependence of the RV variations on the cross-correlation masks constructed for different spectral types (Bouchy et al. 2009).

The Keplerian fit of the RVs was performed simultaneously with the transit modeling (see Sect. 4). All the parameters of the fit are listed in Table 3.

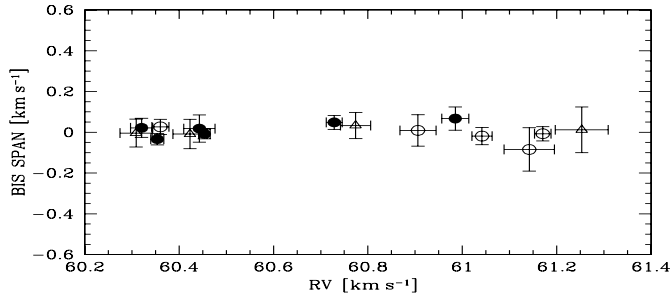
### 4. System parameters

We calculated the flux contamination from nearby stars whose light might fall inside the CoRoT photometric mask using the same methodology as described in Bordé et al. (2010). The method takes into account the photometric mask used to perform the on-board photometry and all the stars in the target neighborhood, including faint background stars. We found that the contamination is less than 0.6% and we further neglected it.

Three sections of the light curve, each centered on a transit, were locally normalized by fitting a third-degree polynomial. Each section was a five-hour interval before the transit ingress and after its egress. The detailed physical modeling of the system was performed by carrying out the transit modeling



**Fig. 3.** The phase-folded radial velocity measurements of CoRoT-20. The various symbols correspond to the different spectrographs used for the follow-up campaign: HARPS (black circle), SOPHIE (open circle) and FIES (open triangle). The best-fit solution is overplotted in full line.



**Fig. 4.** Bisector span versus radial velocity of CoRoT-20 showing no correlation.

and the Keplerian fit of the radial velocity measurements simultaneously. For the transit fit, we used the formalism of Giménez (2006, 2009). The fit involves twelve free parameters: the orbital period  $P$ , the transit epoch  $T_{tr}$ , the transit duration  $T_{14}$ , the ratio of the planet to stellar radii  $R_p/R_*$ , the inclination  $i$  between the orbital plane and the plane of the sky, the Lagrangian orbital elements  $h = e \sin \omega$  and  $k = e \cos \omega$ , where  $e$  is the eccentricity and  $\omega$  the argument of the periastron, the radial velocity semi-amplitude  $K$ , the systemic velocity  $\gamma_{rel}$  and the two offsets between SOPHIE and HARPS radial velocities and SOPHIE and FIES. For the transit modeling, we used a limb-darkening quadratic law (Claret 2003, 2004). The limb-darkening coefficients  $u_a$  and  $u_b$  were taken using the tabulated values for the CoRoT bandpass from Sing (2010) for the atmospheric parameters  $T_{eff}$ ,  $\log g$ , and metallicity derived for the central star (see Sect. 5):  $u_a = 0.4262 \pm 0.0168$  and  $u_b = 0.2434 \pm 0.0108$ . The two corresponding non-linear limb-darkening coefficients are  $u_+ = u_a + u_b = 0.6696 \pm 0.0201$  and  $u_- = u_a - u_b = 0.1828 \pm 0.0201$ . We decided to keep these limb-darkening parameter values fixed in the transit fitting.

The fit was performed using the algorithm AMOEBA (Press et al. 1992). The initial values of the fitted parameters were changed using a Monte Carlo method to find the global minimum of the  $\chi^2$ . The associated  $1-\sigma$  errors were then estimated using a bootstrap procedure described in details in Bouchy et al. (2011). In such a procedure, the limb-darkening parameters were allowed to vary within their error bars related to the atmospheric parameter uncertainties. The final values of the fitted parameters and the subsequently derived system parameters are given

**Table 3.** Planet and star parameters.

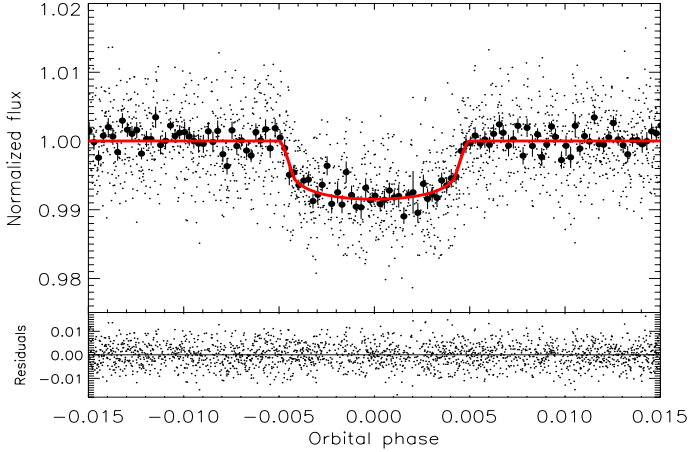
<i>Ephemeris</i>	
Planet orbital period $P$ [days]	$9.24285 \pm 0.00030$
Primary transit epoch $T_{tr}$ [BJD-2 400 000]	$55\,266.0001 \pm 0.0014$
Primary transit duration $T_{14}$ [d]	$0.0927 \pm 0.0019$
Secondary transit epoch $T_s$ [BJD-2 400 000]	$55\,272.46 \pm 0.13$
<i>System parameters</i>	
Periastron epoch $T_{peri}$ [BJD-2 400 000]	$55\,265.79074$
$e \sin \omega$	$0.468 \pm 0.017$
$e \cos \omega$	$0.312 \pm 0.022$
Orbital eccentricity $e$	$0.562 \pm 0.013$
Argument of periastron $\omega$ [deg]	$56.3^{+2.4}_{-2.3}$
Radial velocity semi-amplitude $K$ [m s <sup>-1</sup> ]	$454 \pm 9$
Systemic velocity $\gamma_{rel}$ [km s <sup>-1</sup> ]	$60.623 \pm 0.006$
HARPS-SOPHIE offset velocity $V_{r1}$ [m s <sup>-1</sup> ]	$93 \pm 11$
SOPHIE-FIES offset velocity $V_{r2}$ [m s <sup>-1</sup> ]	$163 \pm 20$
O-C residuals [m s <sup>-1</sup> ]	$27$
Radius ratio $R_p/R_*$	$0.0842 \pm 0.0017$
Impact parameter <sup>a</sup> $b$	$0.26 \pm 0.08$
Scaled semi-major axis $a/R_*$ <sup>b</sup>	$18.95^{+0.63}_{-0.73}$
$M_*^{1/3}/R_*$ [solar units]	$1.022^{+0.034}_{-0.039}$
Stellar density $\rho_*$ [g cm <sup>-3</sup> ]	$1.51^{+0.15}_{-0.17}$
Inclination $i$ [deg]	$88.21 \pm 0.53$
<i>Spectroscopic parameters</i>	
Effective temperature $T_{eff}$ [K]	$5880 \pm 90$
Surface gravity $\log g$ [dex]	$4.20 \pm 0.15$
Metallicity [Fe/H] [dex]	$0.14 \pm 0.12$
Stellar rotational velocity $v \sin i$ [km s <sup>-1</sup> ]	$4.5 \pm 1.0$
Spectral type	G2 V
<i>Stellar and planetary physical parameters from combined analysis</i>	
Star mass [ $M_\odot$ ]	$1.14 \pm 0.08$
Star radius [ $R_\odot$ ]	$1.02 \pm 0.05$
Distance of the system [kpc]	$1.23 \pm 120$
Age of the star $t$ [Myr]	$100^{+800}_{-40}$
Orbital semi-major axis $a$ [AU]	$0.0902 \pm 0.0021$
Orbital distance at periastron $a_{per}$ [AU]	$0.0392 \pm 0.0017$
Orbital distance at apastron $a_{apo}$ [AU]	$0.1409 \pm 0.0037$
Planet mass $M_p$ [ $M_J$ ] <sup>c</sup>	$4.24 \pm 0.23$
Planet radius $R_p$ [ $R_J$ ] <sup>c</sup>	$0.84 \pm 0.04$
Planet density $\rho_p$ [g cm <sup>-3</sup> ]	$8.87 \pm 1.10$
Equilibrium temperature <sup>d</sup> $T_{eq}$ [K]	$1002 \pm 24$
Equilibrium temperature at periastron <sup>d</sup> $T_{eq}^{per}$ [K]	$1444^{+53}_{-46}$
Equilibrium temperature at apastron <sup>d</sup> $T_{eq}^{apa}$ [K]	$764 \pm 18$

**Notes.** (a)  $a/R_* = \frac{1+e \cos \nu_1}{1-e^2} \frac{1+k}{\sqrt{1-\cos^2(\nu_1+\omega-\frac{\pi}{2})\sin^2 i}}$ , where  $\nu_1$  is the true anomaly measured from the periastron passage at the first contact (see Giménez 2009). (b)  $b = \frac{a \cos i}{R_*} \frac{1-e^2}{1+e \sin \omega}$ . (c) Radius and mass of Jupiter taken as 71 492 km and  $1.8986 \times 10^{30}$  g, respectively. (d) Zero albedo equilibrium temperature for an isotropic planetary emission.

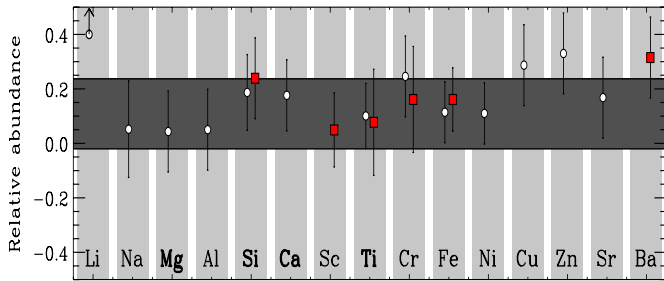
in Table 3. Figure 5 displays the best fit and the observed folded transit.

## 5. Stellar parameters

The spectroscopic analysis was performed in a way similar to that for all CoRoT planets: a master spectrum was created from the co-addition of spectra collected for the radial velocity measurements of the companion. We chose the HARPS spectra that have the highest spectral resolution. We selected those that were unaffected by the Moon reflected light at the time of the observations. Each order of the selected spectra was corrected by the blaze, set in the barycentric rest frame and rebinned to the same



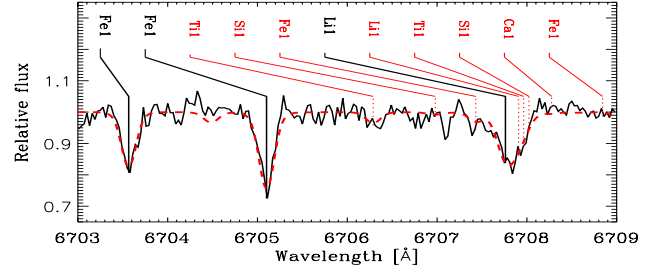
**Fig. 5.** The phase-folded transit in the phase space. The phase bins are 3.3 min and the error bar in each individual bin was calculated as the dispersion of the points inside the bin, divided by the square root of the number of points per bin. The best model is overplotted as a full line.



**Fig. 6.** Abundances of the chemical elements measured with VWA in the HARPS coadded spectrum of CoRoT-20. The abundances refer to the solar value. White circles correspond to neutral lines, red boxes to singly ionized lines, and the yellow area represents the mean metallicity within  $1\text{-}\sigma$  error bar.

wavelength grid with a constant step of  $0.01\text{ \AA}$ . The spectra were then co-added order per order. Each order of the co-added spectrum was then carefully normalized and the overlapping orders were merged resulting in a single one-dimensional spectrum. This master spectrum has a signal-to-noise ratio of 176 per element of resolution at  $5760\text{ \AA}$  in the continuum.

A prior estimate of the atmospheric parameters  $T_{\text{eff}}$ ,  $\log g$ , chemical composition, and  $v \sin i$  was performed by fitting the spectrum with a library of synthetic spectra calculated using MARCS stellar atmosphere models, including the  $H\alpha$  Balmer line. The rotational broadening was estimated for a selection of isolated spectral lines fitted by synthetic spectra convolved with various rotational velocities. We found that  $v \sin i = 4.5 \pm 1.0\text{ km s}^{-1}$  and  $v_{\text{macro}} = 3.5 \pm 1.0\text{ km s}^{-1}$ . The detailed analysis was then carried out using the versatile wavelength analysis package (VWA) (Bruntt et al. 2004, 2010b). A first set of weak and isolated lines of Fe I and Fe II was fitted until the derived abundances of Fe minimized the correlation with the equivalent width and the excitation potentials. We found that  $T_{\text{eff}} = 5880 \pm 90\text{ K}$ ,  $\log g = 4.05 \pm 0.17$  and  $v_{\text{mic}} = 1.10 \pm 0.1\text{ km s}^{-1}$ , which corresponds to a G2-type dwarf. The abundances of other elements for which we could find isolated spectral lines were then derived (see Fig. 6). We performed an independent estimate of the surface gravity from the pressure-sensitive lines: the Mg I b lines, the Na I D doublet, and the Ca I at  $6122\text{ \AA}$  and  $6262\text{ \AA}$ . We fitted the spectrum with the aforementioned grid of synthetic



**Fig. 7.** CoRoT-20 spectrum in a spectral region around the Li I lines at  $6708\text{ \AA}$ .

spectra in regions centered on each of the spectral lines of interest. The inferred value of the surface gravity is  $\log g = 4.2 \pm 0.15$ , a value in good agreement with the  $\log g$  derived with VWA obtained from the agreement between the Fe I and Fe II abundances. We thus adopted  $\log g = 4.2$  for the surface gravity.

The mean metallicity of the star was computed to be the mean of metals with more than ten lines in the spectrum, such as Si, Ca, Ti, Fe, and Ni (Fig. 6). This yields a straight mean of  $[M/H] = 0.14 \pm 0.05$ . The error in  $[M/H]$  caused by the uncertainty in  $T_{\text{eff}}$ ,  $\log g$ , and microturbulence is 0.11 dex, which we must add quadratically to get  $[M/H] = 0.14 \pm 0.12$  (Bruntt et al. 2010a).

We also searched for any indicators of age. We found no hint of stellar activity in the Ca II H and K lines. However, the Li I line is clearly detected at  $6708\text{ \AA}$  (see Fig. 7). We measured an equivalent width of  $W_{\text{eq}} = 44\text{ m\AA}$  and determined a lithium abundance of 2.97. Following Sestito & Randich (2005), this leads to an age in the range from 100 Myr to 1 Gyr, depending on the star's initial rotation velocity.

The modeling of the star in the HR diagram was carried out in the  $(T_{\text{eff}}, M_{\star}^{1/3}/R_{\star})$  plane taking the host star's metallicity into account. This resulted in the final estimates of the star's fundamental parameters given in Table 3 of  $M_{\star} = 1.14 \pm 0.08 M_{\odot}$  and  $R_{\star} = 1.02 \pm 0.05 R_{\odot}$ . The inferred surface gravity is  $\log g = 4.47 \pm 0.11$ , in agreement within the errors with the spectroscopic result. The evolutionary status points to a young star likely in the last stages of the pre-MS phase. We found that the most likely isochrone age is  $100_{-40}^{+800}\text{ Myr}$ , a result in good agreement with the Li abundance.

We calculated the distance to the star. We used the parameters of the star that we had derived and its 2-MASS magnitudes to estimate the reddening. We found a color excess  $E(J - K) = 0.18\text{ mag}$  and the absorption  $A_V = 1.04 \pm 0.5\text{ mag}$  using the extinction law of Schlegel et al. (1998). This yields a distance of  $1.23 \pm 0.12\text{ kpc}$ , consistent with the strong interstellar absorption observed in the Na I (D1) and (D2) lines.

## 6. CoRoT-20 system properties

Compared to the sample of known transiting planets, CoRoT-20b is unusual in many respects. With an orbital period of 9.24 days, it joins the group of transiting planets with periods outside the pile-up at 3 days. It is the fifth planet discovered by CoRoT in this period domain, which currently accounts for 25 planets, 9 of which belong to the multi-planet systems: Kepler-9 (Holman et al. 2010), Kepler-10 (Fressin et al. 2011), and Kepler-11 (Lissauer et al. 2011). However, all these Kepler-planets have a mass that is less than  $\sim 0.3 M_{\text{Jup}}$  and could not be directly compared to the giant planet population.

Excluding these planets in multiple systems, for the 17 remaining objects of the sample that do not have a detected companion, 8, that is 47%, have a significant eccentric orbit with  $e$  in the range from 0.15 to 0.9.

Planets with highly eccentric orbits appear to be found preferentially among the high-mass and/or long-period planet population. With a mass of  $4.13 M_{\text{Jup}}$ , which places it at the border of the gap in mass between the regular hot-Jupiter population and the very massive planet one, CoRoT-20b is consistent with this trend. In the mass-period diagram, these planets are clearly separated from the lighter planets with circular orbits (Pont et al. 2011). This dichotomy and in particular the lack of massive close-in planets on circular orbits suggest that tidal evolution should play an important role in the fate of the planet population.

### 6.1. Tidal evolution

Following the approach of Levrard et al. (2009), we checked the stability of CoRoT-20b to tidal dissipation. The authors calculated the ratio of the total angular momentum of a given system  $L_{\text{tot}}$  to the critical angular momentum  $L_{\text{crit}}$  for some transiting systems. According to Hut (1980), tidal equilibrium states exist when the total angular momentum is larger than this critical value  $L_{\text{crit}}$ . However, this equilibrium state might be either stable or unstable, depending on whether the orbital angular momentum  $L_{\text{orb}}$  is more than three times the total spin angular momentum  $L_{\text{spin}}$ , or not. Levrard et al. (2009) demonstrated that for none of the systems apart from HAT-P-2b the stable tidal equilibrium state, which corresponds to  $L_{\text{tot}}/L_{\text{crit}} > 1$ , exists. In addition, the fate of these close-in planets is ultimately a collision with their host star. The study was recently reexamined and extended to more than 60 transiting systems by Matsumura et al. (2010) who reached a similar conclusion, implying that the vast majority of these close-in planets will spiral-in to their host star and be destroyed by tides. Using Eqs. (1) and (2) given by Levrard et al. (2009) that neglect any effect of a possible magnetized stellar wind, we found for CoRoT-20b that

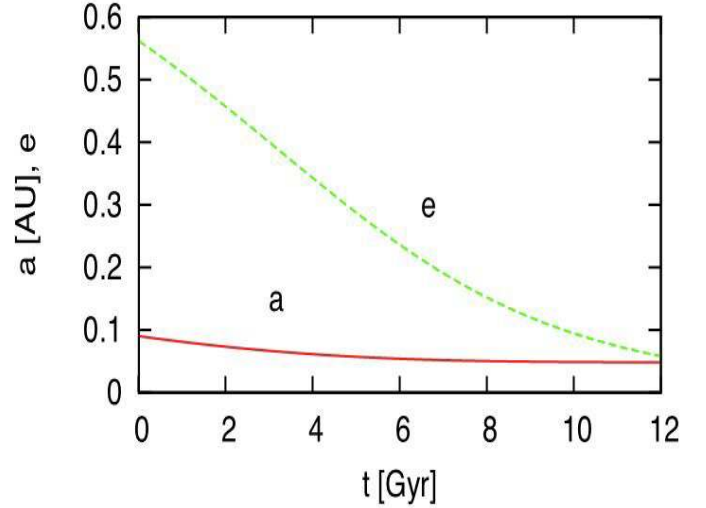
$$L_{\text{tot}}/L_{\text{crit}} = 1.057 \quad \text{and} \quad L_{\text{spin}\star}/L_{\text{orb}} = 0.0458.$$

This shows that, within the current observational uncertainties, the planet has a tidal equilibrium state. We note that our approach also assumes that there is a small value of the stellar obliquity, which is poorly constrained as the star's rotation period could not be derived from the light curve. We simply assumed that the rotation axis is perpendicular to the line of sight and derived the star's rotation period from the values of the  $v \sin i$  (Table 3), which is a regular method for transiting systems. This implies that the rotational period of the star is  $11.5 \pm 3.1$  days, that is of the same order as the planet's orbital period. In the case of CoRoT-20b,  $L_{\text{spin}\star}/L_{\text{orb}} < 1/3$  and most of the angular momentum of the system is in the orbit. According to Matsumura et al. (2010), CoRoT-20b belongs to the very small subgroup of Darwin-stable systems that evolve toward a stable tidal equilibrium state where migration will stop.

From the Roche limit separation, the planet thus lies well beyond two times the Roche limit distance. Using the relation from Faber et al. (2005):

$$a_{\text{R}} = (R_{\text{p}}/0.462)(M_{\star}/M_{\text{p}})^{1/3},$$

we found that the Roche limit  $a_{\text{R}}$  of the system is 0.0057 AU, which supports the migration scenario over the scattering/Kozai-cycle scenario proposed by Ford & Rasio (2006).

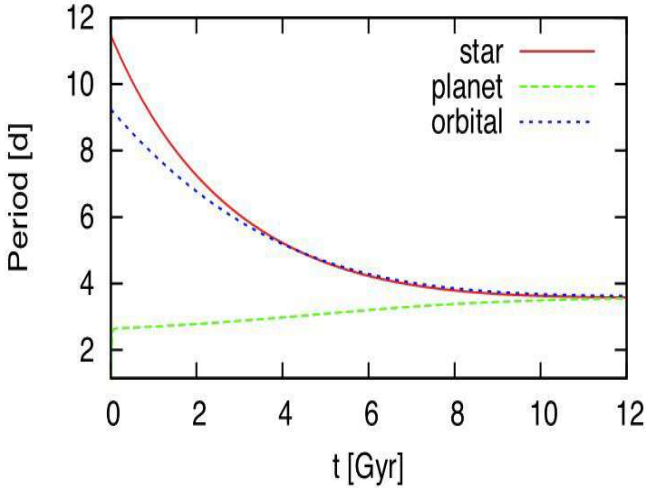


**Fig. 8.** Tidal evolution of the orbital semi-major axis and eccentricity. The figure is displayed on a time interval larger than the expected lifetime of the star to show the triple synchronization characteristic of a Darwin-stable system.

We performed a complete calculation of the tidal evolution of the system formed by the star and the planet assuming a linear tidal model (Mignard 1979; Hut 1981). The main difficulty here is to choose the values of the dissipation in both the star and the planet. For the main semi-diurnal tides of the star, we adopted the value  $Q'_s = 10^7$  as found for hot Jupiters (Hansen 2010; Benítez-Llambay et al. 2011). Owing to the close values of the orbital period and the rotation of the star, the components of the tides raised on the star by the planet that are related to the orbital eccentricity are also equally important, but the values of the current dissipation obtained are of the same order of magnitude. For the planet, we derived one value based on the actually determined  $Q'$  of Jupiter ( $Q' = 1.36 \times 10^5$ , see Lainey et al. 2009). We first note that standard linear tidal theories (see Hut 1981, Eq. (45)) allow us to determine the current rotation period (stationary) of the planet independently of the dissipation. We found  $2.64 \pm 0.13$  days. To transform  $Q'_{\text{Jup}}$  into the planet's  $Q'_p$ , we have to take into account that: (i)  $Q'_p$  scales with the semi-diurnal tide period (see Ferraz-Mello et al. 2008; Matsumura et al. 2010); and (ii)  $Q'_p$  scales with  $R_p^{-5}$  (see Eggleton et al. 1998; Ogilvie & Lin 2004). We found that  $Q'_p = 2.2 \times 10^6$ . Figure 8 shows the variation in the semi-major axis and the evolution of the eccentricity. One can see that with the adopted dissipation values, while the eccentricity tends toward zero, circularization will not be achieved within the lifetime of the system.

Figure 9 shows the evolution of the periods. The planet rotation is currently in a stationary super-synchronous state, that is the planet rotation is faster than its orbital motion. Its period increases as the eccentricity decreases and synchronization is almost reached when the eccentricity becomes very small. The star rotation period is currently decreasing and will equal the synchronous value at some time 4 Gyr from now. However, it will continue to decrease until it reaches the triple synchronous stationary state. The triple synchronization, however, does not seem to be reachable within the lifetime of the star.

We note that the actual  $Q'$  values are not known and the values we used are only estimates found in previous studies. Therefore, the exact timescale of the tidal processes is uncertain. Furthermore, by extracting angular momentum from the system, stellar magnetic braking may prevent the planet from



**Fig. 9.** Tidal evolution of the rotational and orbital periods.

reaching a triple synchronous state and ultimately jeopardize its survival (Bouchy et al. 2011). Simulations in which magnetic braking was active during the whole system lifetime, following the model proposed by Bouvier et al. (1997) and using the same tide parameters as in the examples given above, show that the planet will orbit to within the Roche limit in about 6 Gyr. This result depends critically on the adopted parameters and would require a detailed study that is well beyond the scope of the present paper.

We also investigate the consequences of the circularization of the planet orbit that is in the phase of rapid circularization, on the transits occurrence. Assuming there is no other close massive perturber in the system, two effects appear to cause TTVs: the decrease in the orbital semi-major axis and the circularization of the orbit. The time scale on which the orbital semi-major axis varies is  $\dot{a} = -0.95 \times 10^{-5}$  1/Myr presently (see Fig. 9). As a consequence, a continuous period variation of  $\dot{P}/P \approx -4 \times 10^{-12}$  per cycle is expected. As we know, this linear period variation will cause a parabolic O–C curve, and in 100 years from now the O–C value will be only  $-25$  s. This is slightly over the  $3\sigma$  observation limit of CoRoT (Bean 2009; Csizmadia et al. 2010). Assuming that the transit timing precision can be forced down to 5 s in the future, this O–C value will be reached 45 years from now.

The evolution of the eccentricity during the circularization process has two consequences. First the occurrence of the secondary eclipse will change. The displacement  $D$  of the secondary from phase 0.5 is given by (e.g., Borkovits 2004, Eqs. (1) and (2)). The previous results of the tidal evolution calculations indicated that  $\dot{e} = -4.5 \times 10^{-5}$  1/Myr and  $\dot{P} = -1.5 \times 10^{-3}$  days/Myr. Assuming a constant  $\omega$ , we have that  $\dot{D} = -37.56 \times 10^{-5}$  days/Myr or  $\dot{D} = -9.53 \times 10^{-12}$  days/cycle. This variation is of the same order as the previous one caused by the decreasing semi-major axis, so should be also observable within a century.

For the second effect, that is the circularization of the orbit, one can also consider the occurrence of a small precession in the orbit. This effect is hardly observable, but interesting on the theoretical side, since the transit occurs at the true anomaly  $v = 90^\circ - \omega$ , where  $\omega$  is the argument of periastrion. The latter is also subject to variations caused not only by relativistic effects but also by the tidal effects which force the apsidal line to rotate. However, this variation has a different time scale. We thus

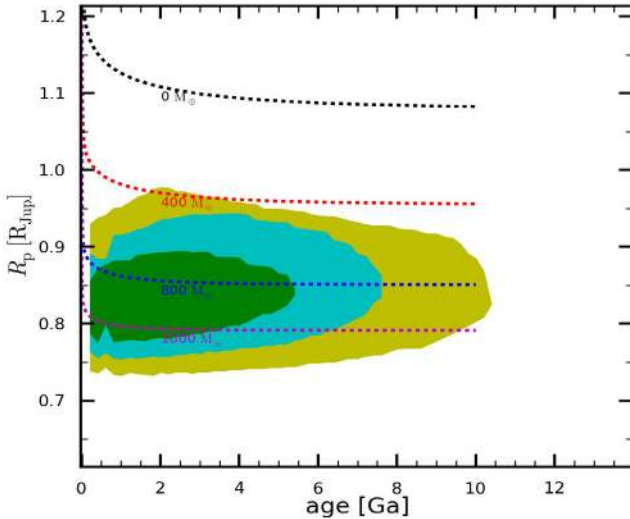
do not take this variation into account here, even if tidal forces also cause a small precession of the orbit showing that  $\dot{\omega}$  is not zero. Hence, if  $e$  decreases due to circularization, and even if  $\omega$  is constant, then at the epoch of transit the eccentric anomaly will increase and hence the mean anomaly at transit will occur later. However, a first estimation shows that this effect may be negligible on a ten year timescale.

## 6.2. Internal structure

The planet CoRoT-20b is a massive hot-Jupiter with a mass of  $4.24 M_{\text{Jup}}$  a radius of  $0.84 M_{\text{Jup}}$ , and an inferred density  $\rho = 8.87 \pm 1.1 \text{ g cm}^{-3}$ . A few giant planets have been already reported that have a similar or even higher density. For example, CoRoT-14b  $\rho = 7.3 \pm 1.5 \text{ g cm}^{-3}$  (Tingley et al. 2011), WASP-18b  $\rho = 8.8 \pm 0.9$  (Hellier et al. 2009; Southworth 2010), or HAT-P-20b  $\rho = 13.78 \pm 1.5 \text{ g cm}^{-3}$  (Bakos et al. 2011). While the masses of these planets spans a large range, from 4 up to more than  $9 M_{\text{Jup}}$ , their radius is close to  $1 R_{\text{Jup}}$ . Given CoRoT-20b’s high planetary mass, its small size is surprising. Among these high density giants planets, only HAT-P-20b is of comparable size, i.e.  $0.867 \pm 0.033 R_{\text{Jup}}$ . The planet CoRoT-20b, as HAT-P-20b, is thus expected to contain large amounts of heavy elements in its interior.

To investigate the internal structure of CoRoT-20b, we computed planetary evolution models with CEPAM (Guillot & Morel 1995), following the description in Guillot & Havel (2011) and Havel et al. (2011) for a planet of a total mass  $4.24 M_{\text{Jup}}$ . We derived a time-averaged equilibrium temperature of the planet to be  $T_{\text{eq}} = 1002 \pm 24$  K. The results for  $T_{\text{eq}} = 1000$  K are shown in Fig. 10 in terms of the planetary size as a function of the system age. The colored regions (green, blue, and yellow) indicate the constraints derived from the CESAM stellar evolution models (Morel & Lebreton 2008) at the 1, 2, and  $3\sigma$  levels, respectively. For preferred ages between 100 Ma and 1 Ga, we find that CoRoT-20b should contain between 680 and 1040  $M_{\oplus}$  of heavy-elements in its interior (i.e. between 50% and 77% of the total planetary mass), at  $1\sigma$  level, about twice the amount needed for HAT-P-20b (see Leconte et al. 2011). While this result is qualitatively in line with the observed correlation between stellar metallicity and heavy elements in the planet (e.g. Guillot et al. 2006; Miller & Fortney 2011, and references therein), the derived amounts are extremely surprising. They would imply that all the heavy elements of a putative gaseous protoplanetary disk of 0.1 to 0.15  $M_{\odot}$  have been filtered out to form CoRoT-20b, hence that an extremely small fraction of hydrogen and helium in that disk was accreted by the planet. This is at odds with the latest accretion formation models (e.g. Ida & Lin 2004; Mordasini et al. 2009). Disk instability with differentiation and partial tidal stripping (Boley & Durisen 2010) is proposed as an alternative formation pathway. According to Boley et al. (2011), this could account for giant planets with massive cores such as HAT-P-20b. In its current orbit, CoRoT-20b does not enter the Roche limit and no tidal stripping is acting on yet but this scenario would deserve further investigations.

We investigated the possibility that changes in the atmospheric model would yield more “reasonable” values for the planetary enrichment. As can be seen from a similar study in the brown dwarf regime (Burrows et al. 2011), the consequences of modified atmospheric properties are limited for objects with the mass of CoRoT-20b (i.e. standard radii for objects of this mass range from 1.05 to 1.20  $R_{\text{Jup}}$ ). By artificially lowering the



**Fig. 10.** Evolution of the size of CoRoT-20b (in Jupiter units) as a function of age (in billion years), compared to constraints inferred from CoRoT photometry, spectroscopy, radial velocimetry and stellar evolution models. Green, blue, and yellow regions correspond to the planetary radii and ages that result from stellar evolution models matching the inferred  $\rho_* - T_{\text{eff}} - [\text{Fe}/\text{H}]$  uncertainty ellipse within  $1\sigma$ ,  $2\sigma$ , and  $3\sigma$ , respectively. Planetary evolution models for a planet with a solar-composition envelope over a central dense core of variable mass ( $0$ ,  $400$ ,  $800$ , and  $1000 M_{\oplus}$  as labelled) are shown as dashed lines. These models also assume that 1% of the incoming stellar irradiation is dissipated deep into the interior of the planet.

infrared atmospheric opacity by a factor 1000 (not shown), we were able to decrease the  $1\sigma$  upper limit to the core mass from  $650$  to  $390 M_{\oplus}$ , a small change compared to the huge and unphysical decrease in the opacity.

On the other hand, one strong assumption in our study is that heavy elements are embedded in a central core. When relatively small amounts of heavy elements are considered, it is not very important whether they are considered as part of a core or mixed in the envelope (e.g. Ikoma et al. 2006). However, as shown by Baraffe et al. (2008), when  $0.5 M_{\text{Jup}}$  of ices are mixed in the envelope of a  $1 M_{\text{Jup}}$  planet, its radius is smaller by  $\sim 0.1 R_{\text{Jup}}$  than when one considers that these elements are part of a central core. It is thus very likely that the mass of heavy elements required to explain the radius of CoRoT-20b is high but significantly smaller than the mass considered here. Our estimates based on the Baraffe et al. (2008) calculations indicate that if mixed in the envelope, a mass of heavy elements two to three times lower than estimated in Fig. 10 would explain the observed planetary size. This would help to explain the formation of the planet, but would still require relatively extreme/unlikely scenarios.

## 7. Summary

We have presented the discovery of CoRoT-20b. The object belongs to the population of massive planets with orbital semi-major axes smaller than  $0.1 \text{ AU}$ , a domain of orbital periods where low and high eccentricity systems co-exist in a narrow range of orbital period. We examined the tidal stability of CoRoT-20 and found that, within the observational uncertainties, it belongs to the relatively small population of transiting planets that are considered as “Darwin-stable”, i.e. systems for which in the absence of processes extracting angular momentum

from the system (i.e. stellar winds), the planet would never fall onto the central star, which would instead be spun-up and achieve triple-synchronization (equality of the orbital, planetary, and stellar spin periods). Other cases in the planetary domain are CoRoT-3b, CoRoT-6b, WASP-7, and HD 80606 (Matsumura et al. 2010). Measuring the stellar obliquity through the Rossiter-McLaughlin effect would provide additional constraints to more tightly constrain its tidal evolution and understand its orbital evolution. The expected semi-amplitude of the radial velocity anomaly of the Rossiter-McLaughlin effect is estimated to be  $22 \pm 5 \text{ m s}^{-1}$ , which is unfortunately quite difficult to detect with the present spectrographs used for the Keplerian orbit determination. Another point would be to assess the presence of additional companions in the system based on long-term radial velocity monitoring as, according to Matsumura et al. (2010), the formation path of close-in planet should be different for single-planet system and multi-planet ones. The CoRoT-20 system thus appears to be an interesting bench test case for studying the tidal orbital and rotational evolution of the close-in population.

The second interesting peculiarity that distinguishes CoRoT-20b from the regular giant planet population is its small observed radius. According to planetary evolution models, the interior of this compact planet should contain a very large amount of heavy elements, with a central dense core whose mass would be in the range between  $680$  and  $1040 M_{\oplus}$ . Although mixing heavy elements in the envelopes rather than confining them to a central core can lead to substantially smaller values (by a factor estimated to be  $\sim 2-3$ ), the origin of such a huge amount of heavy elements is difficult to explain within the framework provided by the current planetary formation models. With HAT-P-20b (Bakos et al. 2011), it is the second example of an extremely metal-rich interior that challenges planetary interior models. However, the two planetary systems differ in many aspects: HAT-P-20b is nearly twice as massive as CoRoT-20b and orbits a K3 metal rich star on a nearly circular orbit, while CoRoT-20 is a solar-type, slightly metal-enriched star. In addition, HAT-P-20 seems to be physically associated with another stellar companion, while up to now, CoRoT-20 has none, either detected or suspected. It thus appears difficult from these two exceptions to derive any trend that would provide information about the origin of these challenging and intriguing bodies.

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