

2007

# Warm molecular hydrogen in the Spitzer SINGS galaxy sample

H Roussel

G Helou

DJ Hollenbach

BT Draine

JD Smith

*See next page for additional authors*

Follow this and additional works at: [https://scholarworks.umass.edu/astro\\_faculty\\_pubs](https://scholarworks.umass.edu/astro_faculty_pubs)



Part of the [Astrophysics and Astronomy Commons](#)

---

## Recommended Citation

Roussel, H; Helou, G; Hollenbach, DJ; Draine, BT; Smith, JD; Armus, L; Schinnerer, E; Walter, F; Engelbracht, CW; Thornley, MD; Kennicutt, RC; Calzetti, D; Dale, DA; Murphy, EJ; and Bot, C, "Warm molecular hydrogen in the Spitzer SINGS galaxy sample" (2007). *ASTROPHYSICAL JOURNAL*. 149.

[10.1086/521667](https://doi.org/10.1086/521667)

This Article is brought to you for free and open access by the Astronomy at ScholarWorks@UMass Amherst. It has been accepted for inclusion in Astronomy Department Faculty Publication Series by an authorized administrator of ScholarWorks@UMass Amherst. For more information, please contact [scholarworks@library.umass.edu](mailto:scholarworks@library.umass.edu).

---

**Authors**

H Roussel, G Helou, DJ Hollenbach, BT Draine, JD Smith, L Armus, E Schinnerer, F Walter, CW Engelbracht, MD Thornley, RC Kennicutt, D Calzetti, DA Dale, EJ Murphy, and C Bot

# Warm molecular hydrogen in the Spitzer SINGS galaxy sample

H. Roussel<sup>1</sup>, G. Helou<sup>2</sup>, D.J. Hollenbach<sup>3</sup>, B.T. Draine<sup>4</sup>, J.D. Smith<sup>5</sup>, L. Armus<sup>2</sup>, E. Schinnerer<sup>1</sup>, F. Walter<sup>1</sup>, C.W. Engelbracht<sup>5</sup>, M.D. Thornley<sup>6</sup>, R.C. Kennicutt<sup>7,5</sup>, D. Calzetti<sup>8</sup>, D.A. Dale<sup>9</sup>, E.J. Murphy<sup>10</sup>, C. Bot<sup>2</sup>

roussel@mpia-hd.mpg.de

## ABSTRACT

Results on the properties of warm molecular hydrogen in 57 normal galaxies are derived from measurements of H<sub>2</sub> rotational transitions in the mid-infrared, obtained as part of the Spitzer Infrared Nearby Galaxies Survey (SINGS). This study extends previous extragalactic surveys of emission lines of H<sub>2</sub>, the most abundant constituent of the molecular interstellar medium, to fainter and more common systems ( $L_{\text{FIR}} = 10^7$  to  $6 \times 10^{10} L_{\odot}$ ) of all morphological and nuclear types. In our sensitive integral-field observations covering kiloparsec-scale areas, the  $17 \mu\text{m}$  S(1) transition is securely detected in the nuclear regions of 86% of SINGS galaxies with stellar masses above  $10^{9.5} M_{\odot}$ . The derived column densities of warm H<sub>2</sub> (with temperatures above  $\sim 100$  K), even though averaged over large areas, are commensurate with values observed in resolved photodissociation regions; the median of the sample is  $3 \times 10^{20} \text{cm}^{-2}$ . They amount to a significant fraction of the column densities of total molecular hydrogen, between 1% and more than 30%. The power emitted in the sum of the three lowest-energy transitions is on average 30% of the power emitted in the bright [SiII] cooling

---

<sup>1</sup>Max-Planck-Institut für Astronomie, Heidelberg, 69117 Germany

<sup>2</sup>California Institute of Technology, Pasadena, CA 91125

<sup>3</sup>NASA Ames Research Center, Moffett Field, CA 94035

<sup>4</sup>Princeton University, Princeton, NJ 08544

<sup>5</sup>Steward Observatory, University of Arizona, Tucson, AZ 85721

<sup>6</sup>Bucknell University, Lewisburg, PA 17837

<sup>7</sup>IoA, University of Cambridge, UK

<sup>8</sup>University of Massachusetts, Amherst, MA 01003

<sup>9</sup>University of Wyoming, Laramie, WY 82071

<sup>10</sup>Yale University, New Haven, CT 06520

line ( $34.8\ \mu\text{m}$ ), and represents about  $4 \times 10^{-4}$  of the total infrared power within the same area for star-forming galaxies, which is consistent with excitation in photodissociation regions. The fact that the  $\text{H}_2$  line intensities scale tightly with the emission in the aromatic bands, even though the average radiation field intensity within the same area varies by a factor ten, can also be understood if both tracers originate predominantly in photodissociation regions, either dense or diffuse. A large fraction of the 25 targets classified as LINERs or Seyferts, however, strongly depart from the rest of the sample, in having warmer  $\text{H}_2$  in the excited states, smaller mass fractions of  $\text{H}_2$  in the warm phase, and an excess of power emitted in the rotational transitions with respect to aromatic bands, the total infrared emission and the [SiII] line. We propose a threshold in  $\text{H}_2$  to aromatic band power ratios, allowing the identification of low-luminosity AGNs by an excess  $\text{H}_2$  excitation. A dominant contribution from shock heating is favored in these objects. Finally, we detect in nearly half the star-forming targets, in particular in low-density central regions of late-type galaxies, non-equilibrium ortho to para ratios, consistent with the effects of pumping by far-ultraviolet photons combined with incomplete ortho-para thermalization by collisions, or possibly non-equilibrium photodissociation fronts advancing into cold gas.

*Subject headings:* galaxies: ISM – infrared: galaxies – infrared: ISM – ISM: lines and bands – ISM: molecules – surveys

## 1. Introduction

Rotational transitions of molecular hydrogen, lying in the mid-infrared range between 5 and  $30\ \mu\text{m}$  (Table 1), provide measurements of the mass and temperature distribution of the bulk of the warm molecular gas phase, at temperatures of  $\sim 100$  to 1000 K. By contrast, the transitions from higher vibrational levels in the near-infrared, better studied because more easily observable from the ground, arise from gas at apparent excitation temperatures of more than 1000 K located in a thin layer of molecular clouds, hosting a negligible fraction (of the order of  $10^{-6}$ ) of the total  $\text{H}_2$  mass (Black & Dalgarno 1976; Burton et al. 1992). The rotational lines are thus more appropriate tracers of molecular gas being exposed to moderate heating, and arise from a much larger volume fraction of the molecular clouds. They constitute one of the most important coolants of warm molecular gas (Neufeld & Kaufman 1993).

$\text{H}_2$  emission lines have been detected in a wide array of sources, including outflows from young stars (Gautier et al. 1976; Bally & Lane 1982), photodissociation regions (Gatley et al.

1987; Tanaka et al. 1989), planetary nebulae (Treffers et al. 1976; Beckwith et al. 1978), supernova remnants (Treffers 1979; Burton et al. 1989), large regions at the centers of galaxies (Thompson et al. 1978; Gatley et al. 1984) and extranuclear large-scale shocks in galaxy collisions (Herbst et al. 1990; Sugai et al. 1997; Appleton et al. 2006). The possible excitation mechanisms are accordingly varied. In normal galaxies, the major excitation source is expected to be the far-ultraviolet radiation of massive stars in photodissociation regions, with photon energies between 6 and 13.6 eV (Hollenbach & Tielens 1997, and references therein). H<sub>2</sub> molecules can be pumped by FUV photons into electronically-excited states, followed by fluorescence and radiative cascade through the vibration-rotation levels of the ground electronic state. Pure fluorescent spectra are produced only if the cascade is not significantly altered by collisions with hydrogen atoms and molecules; if the critical densities for collisional deexcitation are exceeded, a portion of the pump energy is converted to heat by collisions, and the lowest rotational levels are populated by collisions and thermalized. Pure fluorescence is thus much more likely in the vibrational transitions, that have high critical densities, than in the pure rotational transitions considered here, with critical densities below a few 10<sup>3</sup> cm<sup>-3</sup> for S(0) to S(3). Additionally, FUV photons can be absorbed by dust grains, followed by the ejection of photoelectrons that heat the gas. This also results in the thermal excitation of the low-energy levels of H<sub>2</sub> by collisions with the warm gas. Besides the radiation of massive stars, a second important source of excitation is shocks, in molecular outflows, supernova remnants or cloud collisions in a disturbed gravitational potential (Shull & Hollenbach 1978; Draine et al. 1983). In addition to the above processes, X-rays produced in active nuclei or in supernova remnant shocks can partially ionize and heat the gas over large column densities, leading to H<sub>2</sub> excitation by collisions with hydrogen atoms and molecules, and with fast electrons (Lepp & McCray 1983; Draine & Woods 1991; Maloney et al. 1996). Finally, H<sub>2</sub> molecules can be formed directly into excited states.

Surveys of molecular hydrogen line emission in galaxies have been so far mostly restricted to starbursts, active galactic nuclei and ultraluminous systems, and have been performed mostly in the near-infrared, targetting vibration-rotation lines that arise from upper levels with much higher excitation energies than the mid-infrared lines. It has been speculated that the major source of H<sub>2</sub> heating in star-forming galactic nuclei was shocks in supernova remnants, based on comparison of the luminosity of some vibration-rotation H<sub>2</sub> lines with a limited number of Galactic templates and with shock models (Moorwood & Oliva 1988; Mouri et al. 1990). However, scaling individual templates to the integrated emission of galaxies has large inherent uncertainties, and the near-infrared line ratios most often used to discriminate between thermal and non-thermal emission are not always sufficient to distinguish between shocks and fluorescent excitation followed by collisional deexcitation in high-density regions (Sternberg & Dalgarno 1989). Puxley et al. (1988) surveyed star-

burst galaxies in several vibration-rotation lines, and found that the dominant excitation mechanism was pumping by the far-ultraviolet radiation of massive stars, rather than collisional excitation. Davies et al. (2003) reached the same conclusion for a small sample of ultraluminous galaxies, in which the first vibrational level is thermalized by high densities in photodissociation regions. Active nuclei (LINERs or Seyferts) can show an excess of H<sub>2</sub> emission relative to hydrogen recombination lines and aromatic bands (e.g. Moorwood & Oliva 1988; Larkin et al. 1998), but the exact nature of the additional source of excitation, namely X-ray excitation, fluorescence induced by a non-thermal ultraviolet continuum, or shocks induced by dynamical perturbations, is often unclear (e.g. Quillen et al. 1999). It is however unlikely that significant H<sub>2</sub> emission could arise from interaction between molecular clouds and jets from Seyfert nuclei (Rotaciuc et al. 1991; Knop et al. 2001).

The detection of a rotational line of H<sub>2</sub> was first reported by Beck et al. (1979) (the S(2) transition at 12.3  $\mu\text{m}$  in Orion) from observations at Las Campanas Observatory. It was soon followed by many more ground-based detections, but the majority of data on the rotational spectrum of H<sub>2</sub> were produced by the SWS instrument on board ISO (e.g. Lutz et al. 2000; Rigopoulou et al. 2002). Furthermore, with previous infrared spectroscopic capabilities, observations of normal galaxies have proven difficult due to sensitivity limitations, so that our current knowledge is mainly extrapolated from studies of very bright objects, maybe not representative of the general galaxy population. The purpose of this paper is thus to extend previous work to fainter systems than formerly accessible, and to characterize directly the generic properties of the warm molecular hydrogen content of normal galaxies. The SINGS sample (Spitzer Infrared Nearby Galaxies Survey; Kennicutt et al. 2003), covering a broad range of infrared luminosities, morphologies and nuclear types, is ideally suited to such a pursuit.

Studies of rotational lines alone, without information on vibrational levels, have very limited diagnostic value concerning the source of excitation, because the low critical densities of the rotational levels make it likely that they will be thermalized most of the time, and thus cannot be used to distinguish between the various heating mechanisms. Because observations of vibration-rotation transitions in the near-infrared are still scarce for normal galaxies, and because they are typically performed in apertures that are not matched to our observations, we did not attempt to include vibrational levels in our analysis. The characterization of excitation mechanisms and physical conditions in the gas would greatly benefit from such information, but would necessitate an additional dedicated survey.

The rotational lines are, however, energetically important and can characterize the temperature and density conditions of a large mass fraction of the interstellar medium in galaxies, i.e. that consisting of warm molecular gas. From a SWS survey of rotational lines

in nearby starburst and Seyfert galaxies, Valentijn et al. (1996) and Rigopoulou et al. (2002) obtained mass fractions of  $\text{H}_2$  in the warm phase of several percent. In ultraluminous galaxies observed with Spitzer, Higdon et al. (2006) derive much lower mass fractions of warm gas, but the fact that the majority of their sample has only upper limits for the S(0) line makes it possible that the temperatures are overestimated (because computed from the S(1) to S(3) lines only, whenever S(0) is undetected) and thus the masses of warm  $\text{H}_2$  underestimated.

This paper presents observations of warm molecular hydrogen in nearby galaxies obtained as part of SINGS (Kennicutt et al. 2003). From these data, we present quantifications of the temperatures and column densities of warm  $\text{H}_2$  encountered in kiloparsec-scale areas, mostly nuclear regions, and a comparison of the power emitted in the rotational lines with those produced by [SiII] at  $34.8\ \mu\text{m}$ , which is the dominant cooling line of normal galaxies in the mid-infrared range, and by dust. We emphasize the different properties of star-forming regions and nuclei classified as LINERs or Seyferts, and discuss their  $\text{H}_2$  excitation mechanisms. The data, analysis methods and observational results are described in Sections 2 to 5.<sup>1</sup> The interpretation of the main findings is presented in Sections 6 and 7, and a summary of the results and conclusions can be found in Section 8.

## 2. Data and measurements

### 2.1. Targets

The SINGS sample (Kennicutt et al. 2003), comprising 75 galaxies, is intended to be a valuable representative set of local galaxies that are not ultraluminous, and whose moderate distances ensure that the properties of the interstellar medium can be studied at relatively small spatial scales (a few hundreds of parsecs at the shortest wavelengths). Numerous sources with mild starbursts or low-luminosity active nuclei are included. Of this sample, we excluded from the present study the objects that were not observed in spectroscopic mode because of their very low brightness (DDO 154, Ho I, M81 dwA, M81 dwB), or containing very little dust and nebular emission within the nuclear area mapped by the high spectral resolution modules (the quiescent ellipticals NGC 584 and NGC 1404, the quiescent dwarf galaxies DDO 53, DDO 165, Ho IX and the asymmetric magellanic galaxies NGC 4236, NGC 5398, NGC 5408 and IC 4710). The dwarf galaxies IC 2574 and NGC 5474 were also rejected because they lack observations in some of the spectroscopic modules. Of the two

---

<sup>1</sup>For easier comparison to future observations and models, ascii flux tables of all the measured quantities are available upon e-mail request.

star-forming dwarf galaxies with several extranuclear pointings, Ho II and NGC 6822, we retained only NGC 6822 here; the regions within Ho II are indeed too faint to allow an analysis of the H<sub>2</sub> excitation diagram, contrary to some regions within NGC 6822. Low-mass galaxies with extranuclear pointings will be discussed elsewhere. NGC 3034 (M 82) was excluded due to the unavailability of nuclear spectroscopy from SINGS, as well as NGC 1377, which constitutes a galaxy class of its own very different from the rest of the SINGS sample, and has been discussed separately (Roussel et al. 2006). The sample for H<sub>2</sub> measurements comprises 66 targets in 57 galaxies (Table 2). The pointings are centered either on the nuclear regions (for most targets) or on some bright star-forming complexes (for a few dwarf galaxies and a spiral galaxy). Diffuse regions within galactic disks are not covered by the present study.

The aperture over which we extracted the spectra is the intersection of the various areas covered by all four spectroscopic modules. The central position and solid angle of this aperture, used to measure all the quantities presented in this paper (line and continuum fluxes), is listed for each galaxy in Table 2. In practice, the limiting size is that of the maps performed with the high-resolution modules, which were enlarged in a few cases in order to cover the emission from a star-forming circumnuclear ring. At the distances of the targets, the equivalent linear diameters of the apertures range from 60 pc to 3.8 kpc (distribution shown in Fig. 1), and the median is 900 pc. Although the apertures are in general small fractions of the optical extent of the galaxies, the measurements are still averages over very large and complex areas. It is expected that a large number of disconnected star formation sites, in addition to the nucleus, contribute to the total emission.

## 2.2. Broadband imaging

To estimate flux densities of the dust continuum and of the aromatic bands (also referred to as the emission from PAHs, or polycyclic aromatic hydrocarbons), we used images in the 3.6  $\mu\text{m}$  and 7.9  $\mu\text{m}$  bands of the IRAC camera (Fazio et al. 2004), and scan maps in the three bands of the MIPS instrument (Rieke et al. 2004) at effective wavelengths of 24, 71 and 156  $\mu\text{m}$ . Since in early-type galaxies photospheric emission can make an important contribution to 7.9  $\mu\text{m}$  fluxes, we subtracted an estimate of this component in order to obtain pure measurements of aromatic band emission. To this effect, we scaled 3.6  $\mu\text{m}$  fluxes, assumed to be dominated by stellar emission, as described in Helou et al. (2004). The resulting flux densities are noted  $F_{7.9\text{ dust}}$ .

The observing strategy and data reduction are described by Kennicutt et al. (2003). The full width at half maximum of the point spread function (PSF) is close to 2'' at 7.9  $\mu\text{m}$ , 6'' at 24  $\mu\text{m}$ , 18'' at 71  $\mu\text{m}$  and 40'' at 156  $\mu\text{m}$ . Flux calibration uncertainties are of the



order of 10% in the IRAC bands, and 5%, 10% and 15% in the MIPS 24, 71 and 156  $\mu\text{m}$  bands, respectively. To correct for the effects of light scattering in IRAC arrays, we applied to flux densities measured from IRAC maps corrective factors that are appropriate for the photometry of extended sources within apertures of arbitrary size (derived by T. Jarrett and published by Dale et al. 2007). For our apertures, the correction factor at 7.9  $\mu\text{m}$  is of the order of 10%.

### 2.3. Spectroscopic data

The targets were observed in mapping mode with the IRS instrument (Houck et al. 2004), at low spectral resolution between 5 and 38  $\mu\text{m}$ , with the SL and LL slits ( $\lambda/\Delta\lambda \approx 60$ –130) and at high spectral resolution between 10 and 37  $\mu\text{m}$ , with the SH and LH slits ( $\lambda/\Delta\lambda \approx 600$ ). The observing strategy is described by Kennicutt et al. (2003) and Smith et al. (2004). The data were pre-processed with the S13 version of the Spitzer Science Center pipeline. Pixels with an abnormal responsivity were masked, and spectral cubes were built with the Cubism software (Smith et al. 2007a). The flux calibration was performed as described by Roussel et al. (2006). We checked the accuracy of this procedure by systematically comparing broadband fluxes from imaging observations and from spectra, and line fluxes from high and low spectral resolution spectra, for bright lines that are minimally contaminated by broad aromatic features at low resolution (but note that even if the flux calibrations of the different modules were in perfect agreement, deviations would be expected from slight misalignment between the apertures). We obtain  $F_{24}(\text{MIPS})/F_{24}(\text{LL}) = 1.01 \pm 0.04$  (for targets with  $F_{24} > 0.025$  Jy within a diameter of about 50"),  $F_{7.9}(\text{IRAC})/F_{7.9}(\text{SL}) = 0.99 \pm 0.05$  (for targets with  $F_{7.9} > 0.025$  Jy within a diameter of about 30" and accurately determined backgrounds in SL maps),  $F_{[\text{SIII}]}(\text{LL})/F_{[\text{SIII}]}(\text{LH}) = 0.96 \pm 0.14$ ,  $F_{[\text{SIII}] 34}(\text{LL})/F_{[\text{SIII}] 34}(\text{LH}) = 0.92 \pm 0.20$  and  $F_{[\text{SIII}] 19}(\text{LL})/F_{[\text{SIII}] 19}(\text{SH}) = 1.05 \pm 0.21$  (for targets with line fluxes above  $6\sigma$ ).

### 2.4. Measurements

The S(0) to S(3) rotational transitions of  $\text{H}_2$  (Table 1) were measured for all targets. In addition, we measured the S(4) to S(7) transitions in three galaxies in which these lines are bright enough to become detectable at low spectral resolution (see Table 3).

In high spectral resolution data, we defined errors from fluctuations of the pseudo-continuum, which was fitted as an affine function of wavelength ( $F_\nu = a\lambda + b$ ). In SL data, errors at each wavelength were estimated from spatial fluctuations of blank fields within the

satellite spectral maps that are automatically obtained when the source lies in the other half of the slit (see Smith et al. 2004). Both the fluxes and the errors presented in Table 3 were then added linearly for each point of the line profile above the pseudo-continuum. The line profiles were constrained to have a width compatible with the spectral resolution, since the latter is sufficiently low that no line is resolved. Fig. 2 shows the line spectra for the representative galaxies NGC 1097, NGC 6946, NGC 7552, NGC 1266, NGC 4569 and NGC 4579.

The S(1) line is usually the brightest. Of the non-dwarf galaxies of the SINGS sample (with stellar masses estimated as by Lee et al. (2006) above  $10^{9.5} M_{\odot}$ ), the nuclear regions of 86% are securely detected in the S(1) line, with fluxes above three times the measured error. The other 14% are either ellipticals of the LINER type, or late-type spirals (Sc-Sd).

There are two galaxies in common between this sample and that of Rigopoulou et al. (2002), namely NGC 7552 and NGC 6946, the latter from the study of Valentijn et al. (1996). For both, our aperture is larger than the beam of ISO-SWS, which covered an area of 280 to 380 arcsec<sup>2</sup>. For the lines that were detected with SWS, we obtain fluxes that are higher by factors of 2.3 (S(1) in NGC 6946), 5.6 (S(0) in NGC 6946) and 1.1 (S(1) and S(3) in NGC 7552). The exact placement of the ISO-SWS beam is not known. For NGC 6946, given this uncertainty, it is conceivable that the H<sub>2</sub> emission be twice as bright in our 800 arcsec<sup>2</sup> aperture as in the SWS aperture; but the S(0) line flux of Valentijn et al. (1996) is inconsistent with our data.

For this study, we estimate total infrared fluxes (TIR) between 3 and 1100  $\mu\text{m}$ , defined as a linear combination of 24, 71 and 156  $\mu\text{m}$  flux densities. The formula of Dale & Helou (2002) is used here, and we have checked that replacing it with the more recent prescription by Draine & Li (2007) does not change the following results in any appreciable way. The infrared fluxes are measured within the same area as the other quantities for direct comparison. The PSF width at 156  $\mu\text{m}$  is however much larger than the size of our spectroscopic aperture, so that some extrapolation is needed. We first measure MIPS fluxes within the larger aperture used to compare total infrared fluxes with line fluxes measured in the LL module. Then, we scale these fluxes by the ratio of  $F_{24}$  measured in the small aperture to  $F_{24}$  measured in the larger aperture, which is equivalent to assuming that the spectral energy distribution does not change from an area of  $\approx 300$  arcsec<sup>2</sup> to an area of  $\approx 2000$  arcsec<sup>2</sup>. The associated errors are however expected to be small compared with the dynamic range of the quantities discussed in Section 5. Simulations of the overestimation of the far-infrared fluxes caused by the extrapolation, using a simple model of a point-source starburst (with the spectral energy distribution of Mrk 33) superposed on quasi-uniform emission from low radiation field intensity regions (with the colors of the central regions of NGC 24 or NGC 2403), indicate that

the effect should be in most cases of the order of 20% (when the starburst and quiescent components contribute equally at  $156 \mu\text{m}$ ), and in extreme cases reach a maximum of a factor 2 (when the quiescent component dominates). Smith et al. (2007b) reached a similar conclusion (see their Section 3.2).

### 3. Excitation diagrams

Excitation diagrams provide a convenient visualization of the distribution of level populations and allow first constraints on the excitation mechanisms (thermal or non-thermal) that can produce this distribution. They represent the column density in the upper level of each observed transition  $N_u$ , normalized by its statistical weight  $g_u$ , as a function of the upper level energy  $E_u$ . The flux of a transition can be written as  $F = h\nu A N_u \Omega / (4\pi)$ , where  $A$  is the spontaneous emission probability,  $h\nu$  is the transition energy and  $\Omega$  is the beam solid angle. In the assumption of local thermodynamic equilibrium, the total column density  $N_{\text{tot}}$  can be derived from  $N_u = g_u N_{\text{tot}} \exp(-E_u / (kT)) / Z(T)$ , where  $g_u = (2I + 1)(2J + 1)$  is the statistical weight (with the spin number  $I = 0$  for even  $J$  or para transitions, and  $I = 1$  for odd  $J$  or ortho transitions), and  $Z(T) \sim 0.0247 T / (1 - \exp(-6000 \text{ K} / T))$  is the partition function (Herbst et al. 1996), valid for  $T > 40 \text{ K}$ .

The apparent excitation temperature can then be derived from each pair of transitions by:

$$kT = (E_{u2} - E_{u1}) / \ln(N_{u1}/N_{u2} \times g_{u2}/g_{u1}) \quad (1)$$

with  $N_{u1}/N_{u2} = F_1/F_2 \times A_2/A_1 \times \lambda_1/\lambda_2$ . Since both radiative decay and collisions with  $\text{H}_2$  change the rotational number  $J$  by an even number, the ortho and para states are largely decoupled and should in principle be dealt with independently.

#### 3.1. Ortho-para thermalization and departures therefrom

As emphasized by Burton et al. (1992), the lower rotational levels of  $\text{H}_2$  will be in collisional equilibrium over a wide range of conditions, because their critical densities are low. Figure 3 shows the critical densities of all the rotational transitions observable with the IRS instrument, as a function of temperature, computed using the functional form for the collisional de-excitation rate coefficient by  $\text{H}_2$  given by Shull & Beckwith (1982) and the transition probabilities given by Black & Dalgarno (1976). The derived critical densities for each line are about an order of magnitude lower than those for collisions with H computed by Mandy & Martin (1993), the comparison being made at 600 K, since Mandy & Martin

(1993) provide results only for high temperatures.

The integrated emission from warm H<sub>2</sub> in star-forming galaxies is likely to come predominantly from the densest photodissociation regions (PDRs) within the beam, with densities above 10<sup>3</sup> cm<sup>-3</sup> (Burton et al. 1992; Kaufman et al. 2006), in which case the lowest rotational levels will be thermalized. Observations of starburst galaxies with ISO-SWS (Rigopoulou et al. 2002) as well as ultraluminous galaxies with Spitzer-IRS (Higdon et al. 2006) are indeed consistent with this expectation. At first sight, the same applies to the galaxies studied here.

However, some of the excitation diagrams show departures from thermalization of ortho levels with para levels, in the sense that the apparent temperatures derived from each pair of transitions of consecutive rotational number are not monotonic as a function of upper level energy. Clear examples are NGC 1266 ( $T(S0 - S1) = (201 \pm 45)$  K,  $T(S1 - S2) = (465 \pm 34)$  K and  $T(S2 - S3) = (347 \pm 18)$  K); NGC 4254 ( $(162 \pm 9)$  K,  $(358 \pm 59)$  K and  $(259 \pm 38)$  K); and NGC 4631 ( $(127 \pm 8)$  K,  $(342 \pm 39)$  K and  $(268 \pm 25)$  K). Such deviations from thermalization can be explained by an ortho to para density ratio in the excited states apparently different from the equilibrium value. We have

$$\begin{aligned} OPR &= \frac{OPR_{\text{high T}}}{3} \frac{\sum_o (2I_o+1) (2J_o+1) \exp(-E_o / (kT))}{\sum_p (2I_p+1) (2J_p+1) \exp(-E_p / (kT))} \\ &= OPR_{\text{high T}} \frac{\sum_o (2J_o+1) \exp(-E_o / (kT))}{\sum_p (2J_p+1) \exp(-E_p / (kT))} \end{aligned} \quad (2)$$

where the subscripts *o* and *p* designate ortho and para levels respectively ( $I_p = 0$  and  $I_o = 1$ ).  $OPR_{\text{high T}}$ , equal to the actual ortho to para ratio ( $OPR$ ) in the high-temperature limit, expresses deviations from local thermodynamic equilibrium (LTE) if it differs from three. It may be called the effective nuclear spin degeneracy ratio, but will hereafter be called the ortho to para ratio for convenience. In LTE,  $OPR \sim 2$  for  $T \sim 100$  K and  $OPR \sim 3$  for  $T > 200$  K (Burton et al. 1992), but  $OPR_{\text{high T}} = 3$  at all temperatures. Although  $OPR_{\text{high T}} < 3$  may be inferred for the excited states ( $J \geq 2$ ), this does not imply that the ortho to para ratio of the bulk of the gas in the  $J = 1$  and  $J = 0$  states be out of LTE. In the following, LTE will refer more particularly to the equilibrium between the ortho and para levels, and not of the ortho levels or para levels separately. Extinction effects are discussed in Section 3.3 and the interpretation of  $OPR_{\text{high T}}$  values is postponed to Section 6.

To derive temperatures and column densities, we first determine whether the excitation diagram is compatible or not with LTE by inserting explicitly the factor  $OPR_{\text{high T}}/3$  in the equations for column densities of the ortho levels, and deriving temperatures from each pair of consecutive transitions as a function of  $OPR_{\text{high T}}$ , to verify whether these conditions are satisfied:  $T(S0 - S1) \leq T(S0 - S2) \leq T(S1 - S2) \leq T(S1 - S3) \leq T(S2 - S3)$ , since in gas with a distribution of temperatures, ratios of transitions with low-energy upper levels

always probe lower excitation temperatures than ratios of transitions with higher-energy upper levels.  $T(S0 - S2)$  and  $T(S1 - S3)$  are independent of  $OPR_{\text{high T}}$  and determined directly from the observed fluxes, but  $T(S0 - S1)$ ,  $T(S1 - S2)$  and  $T(S2 - S3)$  depend on  $OPR_{\text{high T}}$ . For each pair  $(p, o) = (0,1)$ ,  $(2,1)$  and  $(2,3)$ , we have:

$$kT(S_p - S_o) = (E_{u_o} - E_{u_p}) / \ln(OPR_{\text{high T}} \times R). \quad (3)$$

with  $R = F_p/F_o \times A_o/A_p \times \lambda_p/\lambda_o \times (2J_o + 1) / (2J_p + 1)$ . Figure 4 shows the corresponding diagram for two galaxies. In case the above condition on the temperatures is satisfied for  $OPR_{\text{high T}} = 3$ , as illustrated for NGC 3198, we fix  $OPR_{\text{high T}} = 3$ ; in the opposite case, illustrated by NGC 4631, we fit  $OPR_{\text{high T}}$  as explained below. The excitation diagrams of all the galaxies, with fits overlaid, are shown in Fig. 5.

### 3.2. Temperatures and column densities

Since in all cases the excitation diagrams indicate that a single temperature does not fit all the line fluxes, we assume that the  $H_2$  emission is the sum of two discrete components of different temperatures, which is enough to reproduce accurately the observed fluxes. In the general case of  $OPR_{\text{high T}} = 3$ , we perform a least-squares fit of the excitation diagram to determine the parameters of the two discrete components (the lower temperature  $T_1$ , the upper temperature  $T_2$ , and their mass fraction) and the normalization by the total column density. The results are listed in Table 4.

When the gas is at a range of temperatures, it is in practice impossible to lift the degeneracy between mass and temperature from the lowest-energy levels. Since the column density has a very steep dependence on  $T_1$ , we adopt two different procedures to fit the excitation diagrams and ascertain the amplitude of the uncertainties caused by this degeneracy. In the first case, we constrain  $T_1$  to exceed the value for which the column density is 20% higher than the nominal density derived from  $T(S0 - S1)$ . In the second case, we leave  $T_1$  unconstrained. In the following, both approaches will be retained when discussing results that depend on  $T_1$ .

For the results not to be biased by systematic sensitivity differences at the wavelengths of the  $H_2$  transitions, we also replace the measured errors by a uniform weight. When  $OPR_{\text{high T}}$  is allowed to be fitted, we fix  $T_2$  at  $1.3 \times T(S1 - S3)$  in the constrained- $T_1$  fits, which was chosen from the median value of  $T_2$  in galaxies with  $OPR_{\text{high T}} = 3$ . In free- $T_1$  fits with  $OPR_{\text{high T}} = 3$ , the distribution of  $T_2/T(S1 - S3)$  is large, with a tail of high values; therefore,  $T_2$  is first fixed at the median value,  $1.14 \times T(S1 - S3)$ , and then at  $1.5 \times T(S1 - S3)$ , to probe the full range of most likely values. Finally, when one flux is an upper limit, we fix

both  $T_1$  at  $0.98 \times T(S0 - S1)$  (which increases the total column density by a maximum of  $\sim 20\%$  with respect to that obtained with  $T_1 = T(S0 - S1)$  but allows a small contribution from hotter gas to the S(0) and S(1) lines), and  $T_2$  as above.

For the three galaxies from which more transitions, up to S(7), could be measured, the procedure is the same except that a third component has to be added. The additional parameters are  $T_3$  and the mass fraction of the second component, and  $T_2$  is fixed at 400 K.

Several galaxies barely satisfy the criterion on temperatures to have  $OPR_{\text{high T}} = 3$ , with  $T(S1 - S2) \geq 0.95 \times T(S1 - S3)$  and  $T(S2 - S3) \leq 1.05 \times T(S1 - S3)$ . When  $T_1$  is constrained, the quality of their fits can be improved by allowing  $OPR_{\text{high T}}$  to vary. For these objects, we provide results with  $OPR_{\text{high T}} < 3$ . Allowing  $OPR_{\text{high T}}$  to be smaller than the equilibrium value has the indirect consequence that the derived column densities are smaller. The amplitude of this effect is indicated in Table 4. Similarly, for NGC 1705 and NGC 4552, we provide results with  $OPR_{\text{high T}} < 3$  and indicate the change in column density with respect to  $OPR_{\text{high T}} = 3$ , because although the S(2) transition being an upper limit prevents any reliable determination of  $OPR_{\text{high T}}$ , the  $T_1$  temperatures derived with  $OPR_{\text{high T}} = 3$  are the two lowest of the whole sample, raising the suspicion that they might be artifacts of the constraint on  $OPR_{\text{high T}}$ . We also consider  $OPR_{\text{high T}} < 3$  more likely for these galaxies in view of the dependence of  $OPR_{\text{high T}}$  on  $H_2$  brightness, discussed later in Section 6.

The median  $T_1$  temperature is 154 K when the fits are constrained (ranging between 97 and 300 K); when no constraint is applied, the median  $T_1$  is 118 K with  $T_2 = 1.14 \times T(S1 - S3)$ , and 161 K with  $T_2 = 1.5 \times T(S1 - S3)$ . The total column densities that we obtained, averaged over kiloparsec-scale regions in galactic centers, range between  $10^{19}$  and  $2 \times 10^{21} \text{ cm}^{-2}$  (for constrained- $T_1$  fits), or  $2 \times 10^{22} \text{ cm}^{-2}$  (for free- $T_1$  fits), and their medians are respectively  $3 \times 10^{20} \text{ cm}^{-2}$  and  $5\text{--}6 \times 10^{20} \text{ cm}^{-2}$  (Fig. 6). This can be compared with typical column densities of resolved photodissociation regions in the Milky Way. In the Orion Bar, column densities of  $H_2$  warmer than 400 K, derived from rotational lines, lie between  $10^{20}$  and  $10^{21} \text{ cm}^{-2}$  (Parmar et al. 1991; Allers et al. 2005). Note that because the Orion Bar is observed nearly edge-on, an equivalent PDR seen face-on would have lower column densities. In NGC 7023, Fuente et al. (1999) derived a total column density of  $5 \times 10^{20} \text{ cm}^{-2}$  for  $H_2$  warmer than 300 K. Thus, if the  $H_2$  emission in our targets comes from similar photodissociation regions, they must occupy in general a very large fraction of the observing beam, assuming that they do not overlap on the line of sight.

Figure 6 also shows a clear dependence of the local (nuclear) column density of warm  $H_2$  on the total stellar mass of the host galaxy. The stellar mass and the infrared luminosity being correlated for star-forming galaxies, there is a similar dependence on far-infrared luminosities.

To first order, the column density of warm  $\text{H}_2$  shows the same behavior as tracers of molecular gas and star formation rate densities, which suggests that the primary source of  $\text{H}_2$  heating is the star formation activity in non-AGN galaxies, and the nuclear regions respond to the global mass and luminosity. LINER and Sy nuclei do not follow the correlation shown by star-forming regions, and tend to have smaller column densities of warm  $\text{H}_2$ . The differences in terms of energy output and excitation mechanisms will be studied in more detail in Sections 5, 6 and 7. Since the few extranuclear regions and dwarf galaxies included in the sample do not distinguish themselves from the other star-forming targets in any obvious way, here and in the following, they are not discussed as separate categories.

### 3.3. Optical depth toward $\text{H}_2$

Consistent with the negligible optical depths inferred from the silicate absorption bands at  $10\ \mu\text{m}$  and  $18\ \mu\text{m}$  in most SINGS galaxies (Smith et al. 2007b), that support the modest values of nebular extinction derived from the Balmer decrement (Dale et al. 2006), we assume zero extinction both in the lines and in the dust continuum for all the targets. In eight galactic centers among the SINGS sample (included here), Smith et al. (2007b) obtained a better fit in their decomposition of the low spectral resolution spectra by including a finite optical depth in the silicate bands. We expect the warm  $\text{H}_2$  component to suffer less extinction, on average, than the warm dust continuum, because the two emission sources will not be cospatial in general, and the regions of high optical depth will be confined to compact regions, probably more concentrated than the regions participating in  $\text{H}_2$  emission (see the striking example of NGC 1377; Roussel et al. 2006). In particular, Higdon et al. (2006) did not see any evidence for significant extinction in the rotational  $\text{H}_2$  lines of ultraluminous galaxies, although these objects are expected to have much higher optical depths than the present sample. In the absence of any quantitative constraint on the differential extinction between the dust and  $\text{H}_2$ , we do not attempt to correct  $\text{H}_2$  fluxes for extinction.

Using the extinction law of Moneti et al. (2001), valid for the Galactic center, we have  $A(9.7\ \mu\text{m})/A_V = 0.15$ ,  $A(28.2\ \mu\text{m})/A(9.7\ \mu\text{m}) = 0.25$  and  $A(17.0\ \mu\text{m})/A(9.7\ \mu\text{m}) = A(12.3\ \mu\text{m})/A(9.7\ \mu\text{m}) = 0.46$ . Even assuming the same optical depth toward the warm molecular hydrogen as toward the hot dust, the extinction correction would not change significantly the derived column densities. The extinction is modest at  $10\ \mu\text{m}$ , and therefore negligible at  $28\ \mu\text{m}$ , the wavelength of the S(0) line which dominates the total column density determination. Extinction effects would however depress the S(1) and S(3) line fluxes with respect to S(0) and S(2), and could thus artificially lower the derived  $OPR_{\text{high T}}$ . In the following, we put lower limits to  $OPR_{\text{high T}}$  values, when less than 3, derived for the eight

galaxies with non-zero optical depth at  $10\ \mu\text{m}$ .

NGC 3198 is the sample galaxy with the highest optical depth in the silicate feature according to Smith et al. (2007b), but its excitation diagram shows no sign of attenuation of the S(1) and S(3) lines relative to the others, and is consistent with  $OPR_{\text{high T}} = 3$  (Fig. 4). The second most obscured galaxy of the present sample is NGC 1266 (it also has the highest nebular extinction according to Dale et al. (2006),  $A_V = 4.1$  mag), for which we derive  $OPR_{\text{high T}} < 3$ . If this were due to optical depth effects, then the S(3) line at  $9.7\ \mu\text{m}$  should be more attenuated than the S(1) line at  $17.0\ \mu\text{m}$ . Since this would be consistent with the excitation diagram, we cannot exclude that the apparently low  $OPR_{\text{high T}}$  value be an extinction artifact in at least this galaxy. The dissimilar behavior of the two galaxies in terms of differential extinction between  $\text{H}_2$  and the dust could then arise from different excitation mechanisms and geometries: whereas in the nuclear regions of NGC 3198, classified as purely HII, the  $\text{H}_2$  emission is presumably distributed over a large volume, the  $\text{H}_2$  emission in the LINER nucleus of NGC 1266 may be much more compact, and not produced by star formation processes (see Section 5). For 13 galaxies with negligible silicate extinction in the spectral decomposition performed by Smith et al. (2007b), the excitation diagrams do imply  $OPR_{\text{high T}} < 3$ , whether a constraint on the lower temperature  $T_1$  is applied or not. In addition, of the 6 galaxies found to have non-zero silicate extinction and  $OPR_{\text{high T}} < 3$ , three would require  $\tau(\text{H}_2) > \tau_{\text{sil}}$  in order to obtain  $OPR_{\text{high T}} = 3$  after extinction correction (by  $\geq 25\%$  for NGC 1266, by a factor  $\geq 6$  for NGC 4631 and by a factor  $\geq 3.5$  for NGC 5866). The three others (NGC 1482, 4536 and 6946) would require either  $\tau(\text{H}_2) > \tau_{\text{sil}}$ , or very low  $T_1$  temperatures ( $\leq 100$  K). Since it is unlikely that the optical depth toward  $\text{H}_2$  be higher than toward the dust continuum, we conclude that our finding, discussed in Section 6, is robust against extinction effects.

#### 4. Mass fraction in the warm phase

In order to estimate the fraction of molecular hydrogen that is heated to temperatures above  $\sim 100$  K, we searched the literature for observed intensities of the 2.6 mm CO(1-0) line within a beam comparable to the solid angle of our observations. Table 5 summarizes the adopted data. The column density of cold  $\text{H}_2$  as given here is derived from CO velocity-integrated intensities on the main-beam temperature scale, assuming a uniform conversion factor of CO(1-0) intensities to  $\text{H}_2$  column densities of  $2.3 \times 10^{20}\ \text{cm}^{-2}/(\text{K km s}^{-1})$  (Strong et al. 1988). We derived aperture corrections to the CO intensities by projecting on a map both the IRS beam and the CO beam. We did not use any deconvolution technique. Whenever possible, a map from the BIMA SONG interferometric survey, including the zero-



spacing total intensity (Helfer et al. 2003), was used. Otherwise, we used instead the  $7.9\ \mu\text{m}$  map and assumed the spatial distributions of aromatic bands in emission and CO(1-0) line emission to be similar at the large spatial scales corresponding to our apertures. This can be justified qualitatively by the association of dust with molecular gas and the Schmidt law (for a recent study of the spatially-resolved Schmidt law, see Kennicutt et al. 2007, in preparation). The applied correction factors are listed in Table 5. In some cases, there are several available measurements all giving consistent estimates to within 30% ; the corresponding unused references are given within parentheses.

There are two major sources of uncertainty in this comparison. The first one is inherent to the difficulty of matching the physical area covered by the IRS integral-field measurements, from single-dish or aperture-synthesis measurements within a different beam. The second dominant source of uncertainty comes from the conversion factor of CO intensities to  $\text{H}_2$  masses, assumed uniform here. The result of Strong et al. (1988) is derived from a comparison of Galactic  $\gamma$ -ray emission with CO and HI emission. Dame et al. (2001) obtained a consistent conversion factor by extrapolating the gas-to-dust mass ratio measured from HI and far-infrared emission, in areas devoid of CO emission, to molecular clouds. Both methods provide an estimate of the total  $\text{H}_2$  column density, including the warm gas as well as the cold gas, for molecular clouds under similar average physical conditions as Galactic clouds. Note however that conversion factors both significantly lower and significantly higher have been derived for normal galaxies. For instance, the recent study of Draine et al. (2007) favors an average value of  $4 \times 10^{20}\ \text{cm}^{-2}/(\text{K km s}^{-1})$ , based on global gas-to-dust mass ratios in the SINGS sample. In addition, the ratio of  $\text{H}_2$  column density to CO intensity can vary by at least a factor two, depending on the physical conditions of the regions emitting in CO (Maloney & Black 1988), even though our observing aperture is large enough to cover a large number of molecular clouds and dilute some of the dispersion in their physical properties. In particular, the conversion factor is expected to be lower for compact and actively star-forming regions than for more diffuse and more quiescent regions. We discount here variations due to metal abundance, since we could find CO measurements for only two low-metallicity targets (NGC 2915 and NGC 6822\_A).

Figure 7 shows the mass fraction of molecular hydrogen in the warm phase ( $T \geq T_1 \approx 100\ \text{K}$ ) as a function of the minimum temperature of the warm component, as determined by the lowest-energy rotational  $\text{H}_2$  lines. The nuclei classified as star-forming have a relatively narrow range of lower temperatures ( $T_1 = 144 \pm 24\ \text{K}$  for 31 nuclei, with or without CO data, from the constrained fits). However, for nuclear regions classified as LINERs or Seyferts, the spread in temperatures is higher ( $T_1 = 180 \pm 45\ \text{K}$  for 25 nuclei). No statistically-significant difference exists between the 18 LINER and 7 Sy nuclei.

A clear anticorrelation exists between the two quantities plotted (partly the result of the degeneracy between temperatures and column densities), which remains intact when restricting the sample to those galaxies for which we could find well-matched CO data (i.e. with correction factors close to unity and with several consistent measurements). The dynamic range in the warm gas mass fraction is much higher than accounted for by the uncertainty on the total H<sub>2</sub> mass. The uncertainty on the warm H<sub>2</sub> mass for individual objects is however extremely large, owing to the degeneracy between  $T_1$ , often ill-constrained by the data, and the column density. The example of NGC 4579 is the most striking (see Table 4). Since its rotational levels up to J=5 are close to thermal equilibrium (at a single temperature of the order of 300-400 K), such a component at 70 K as found in the free- $T_1$  fit is unlikely to be real. Because the fits where  $T_1$  is unconstrained allow mass fractions in the warm phase that are sometimes unphysical (for example for NGC 2976 and NGC 4826), we favor the constrained fits as more plausible, but emphasize that the mass distribution at low temperatures is in general unconstrained.

In the case of constrained- $T_1$  fits, it appears that for a small set of nuclear regions classified as LINERs or Seyferts, the warm H<sub>2</sub> phase consists only of a very small fraction of the total mass, but heated to higher temperatures than in regions classified as purely star-forming. This behavior arises naturally if normal photodissociation region excitation is missing, and if the hotter gas is located in a thin layer of molecular clouds, or has a small filling factor. In the case of free- $T_1$  fits, only NGC 1316 (Fornax A) remains robustly in the part of the diagram with high  $T_1$  and mass fraction below 3%. The average temperature is however still higher for LINERs and Seyferts than for HII nuclei, and the average mass fraction in the warm phase likewise lower. The reason for this difference will be further discussed in Section 7, addressing the excitation mechanisms.

## 5. Comparison of the powers emitted by warm H<sub>2</sub>, [SiII] and dust in star-forming regions

In order to empirically quantify the importance of the H<sub>2</sub> rotational lines in cooling the interstellar medium of normal galaxies, and to put constraints on the possible excitation mechanisms of H<sub>2</sub>, discussed in more detail in Section 6, we examine power ratios of H<sub>2</sub> to other tracers of the warm interstellar medium extracted from the same observations. The results presented here are independent of any fits to the excitation diagrams. Only the HII nuclei and complexes are considered, LINER and Sy nuclei being separately discussed in Section 7. Since the bulk of warm H<sub>2</sub>, at the lowest rotational temperatures, emits mostly in the S(0) to S(2) lines, whereas the S(3) line emission has a noticeably higher contribution from

hotter  $\text{H}_2$ , probably indicating more mixed excitation sources (anticipating the discussion of excitation mechanisms, see Section 6.1), we choose, as the most useful quantification of  $\text{H}_2$  rotational emission in star-forming targets, the sum of the S(0) to S(2) lines.

### 5.1. Total infrared emission

In photodissociation regions, almost all the far-ultraviolet power from massive stars that does not escape is absorbed by dust and converted to infrared continuum radiation, or is absorbed by  $\text{H}_2$ . Only a very small fraction of the power absorbed by dust, of the order of 1%, is converted to photoelectrons that heat the gas, and emerges as infrared lines (Tielens & Hollenbach 1985). The dominant gas coolants are the [OI] and [CII] lines at  $63\ \mu\text{m}$  and  $158\ \mu\text{m}$ , but mid-infrared lines, in particular [SiII] at  $34.8\ \mu\text{m}$  and the  $\text{H}_2$  rotational lines, are also energetically significant. Although the transition rate coefficients of  $\text{H}_2$  are low and the excitation energies relatively high,  $\text{H}_2$  molecules are dominant in number.

The observed ratios of the power emitted in the sum of the S(0) to S(2) lines to the total dust power emitted in the infrared (TIR; see Section 2.4) range between  $2.5 \times 10^{-4}$  and  $7.5 \times 10^{-4}$  for nuclear regions that are not classified as LINERs or Seyferts (Fig. 8a). These ratios are in agreement with predictions of the photodissociation models of Kaufman et al. (2006) for a wide variety of radiation field intensities  $G_0$  and hydrogen densities  $n$ , but a relatively narrow range of  $G_0/n$  ratios, approximately between 0.1 and 1 with  $G_0$  in units of  $1.6 \times 10^{-3}\ \text{erg s}^{-1}\ \text{cm}^{-2}$  and  $n$  in units of  $\text{cm}^{-3}$ . Note that models predict the ratio of the  $\text{H}_2$  line power to the far-ultraviolet (FUV) power (for photon energies between 6 and 13.6 eV), rather than the total infrared power. Since the intrinsic FUV flux heating the photodissociation regions is unknown, the comparison between observations and models is here made by assuming an exact conversion of FUV photons to infrared photons. The fraction of dust heating provided by non-FUV photons can however be significant. Allowing for this effect would reduce the derived  $G_0/n$  ratios. The  $\text{H}_2$  rotational line fluxes predicted by Kaufman et al. (2006) are nearly an order of magnitude higher than those from the older models of Burton et al. (1992), because of the inclusion of photoelectric heating by PAHs, a better  $\text{H}_2$  model, and a finer numerical grid near the region of  $\text{H}_2$  emission.

The inferred  $G_0/n$  ratios are lower than the results of Malhotra et al. (2001), who derived the physical conditions of an ensemble of bright star-forming galaxies from the [CII] and [OI] lines. They found  $G_0/n$  ratios between about 0.5 and 6, i.e. on average 5 times higher than those indicated here by the rotational  $\text{H}_2$  lines. A possible explanation is that  $\text{H}_2$  emission comes from cooler and denser regions than [CII] and [OI], because  $\text{H}_2$  exists at higher optical depths inside the clouds than  $\text{C}^+$  and O (Hollenbach & Tielens 1997). The dif-

ference in physical conditions could thus merely reflect a different spatial origin. Besides the different locations within PDRs, the two studies also deal with different regions within galaxies: the targets of Malhotra et al. (2001) were selected to have most of their line emission encompassed by the ISO-LWS beam of  $70''$ , whereas our apertures usually cover small fractions of the line and dust emitting areas. Alternatively, the observations of Malhotra et al. (2001) could reflect intrinsically different physical conditions because their sample contains galaxies on average brighter and more active than the sample used here. Their far-infrared luminosities (in the definition of Helou et al. 1988) range from  $6 \times 10^7$  to  $8 \times 10^{11} L_{\odot}$ , with a median of  $1.5 \times 10^{10} L_{\odot}$ , whereas the far-infrared luminosities of the present sample range from  $10^7$  to  $6 \times 10^{10} L_{\odot}$ , with a median of  $3 \times 10^9 L_{\odot}$ . The median  $F_{60}/F_{100}$  ratio is also higher in the sample of Malhotra et al. (2001) (0.57) than in our sample (0.41), indicating higher radiation field intensities on average. The  $G_0/n$  ratios derived by Malhotra et al. (2001) however do not display any clear correlation with either infrared luminosity or color. Only NGC 1482 and NGC 5713, included in both samples, allow a direct comparison of model results (we discard the LINER NGC 1266 because most of its  $H_2$  emission is not produced by PDRs, as shown in Sect. 7). For both sources, the  $H_2$  line fluxes indicate consistently  $G_0 \sim 4000$  and  $n \sim 1-2 \times 10^4$ . For NGC 1482,  $G_0$  is in agreement with one of the two models of Malhotra et al. (2001), but  $n$  is at least four times higher. For NGC 5713,  $G_0$  is two times higher than that of Malhotra et al. (2001), and  $n$  is at least six times higher. In conclusion, we favor differences in spatial origin (both within PDRs and within galaxies) as a likely cause for the different model results.

## 5.2. [SiII] line emission

Figure 8b shows the ratio of powers emitted in the  $H_2$  rotational lines and in the [SiII] line. The dispersion in the ratio is very similar to that seen in Fig. 8a, and the [SiII] line alone emits more power than the sum of the S(0) to S(3) transitions in HII nuclei. The [SiII] line has indeed been found to be the brightest mid-infrared cooling line and to scale tightly with the total infrared power both in nuclear and extranuclear regions within the SINGS sample galaxies (Helou et al., in preparation), with only a very slight dependence on the radiation field intensity. We have on average  $F(S0 - S2)/F([SiII]) = 0.3$  (ranging between 0.15 and 0.5 for nuclei), and  $F([SiII])/TIR = 2 \times 10^{-3}$ . Using the [SiII] line as a substitute for the total dust emission is advantageous because it is observed at about the same angular resolution as the  $H_2$  lines, whereas estimating the total infrared power within these apertures requires a large extrapolation (because of the large width of the point spread function at 70 and  $160 \mu\text{m}$ ), making the uncertainty on  $H_2/TIR$  relatively high. The [SiII] power predicted by the photodissociation region model of Kaufman et al. (2006), with the

same physical conditions as above, is however smaller than observed by a factor greater than 3, which implies either that the majority of [SiII] emission comes from HII regions in high-metallicity nuclear regions, or that the fraction of silicon incorporated in dust grains is smaller than 90%.

Only the regions B and C in NGC6822 have significantly less [SiII] emission, with respect to H<sub>2</sub> emission, than the nuclear regions of spiral galaxies. Their H<sub>2</sub> emission is also slightly overluminous with respect to the aromatic bands (Fig. 8c). This may not be entirely attributable to a metallicity effect, decreasing the abundances of PAHs and silicon, since region A (Hubble V) has normal flux ratios, and oxygen abundances are quite uniform in NGC6822 (Pagel et al. 1980). An alternative explanation is that additional excitation of H<sub>2</sub> may be provided in regions B and C, with respect to region A, by shocks in supernova remnants (see the more general discussion in Sect. 7.2). To our knowledge, no independent evidence exists to test the existence of shocks in these regions. Chandar et al. (2000) obtained a normal HII optical line spectrum at the center of NGC6822\_C, but since their beam of 2.5 arcsec<sup>2</sup> is only about 1% of ours, we cannot rule out shock excitation. Finally, given the small distance of NGC6822, the regions covered by the IRS aperture are less than 100 pc in size. Greater fluctuations around the average properties are thus not unexpected. At present, we are unable to decide which scenario is the most likely.

### 5.3. Aromatic bands

Figure 8c shows a remarkable constancy of the power ratio of the H<sub>2</sub> rotational lines to the aromatic bands. Among the measured dust and gas observables, PAH emission provides the tightest correlation with H<sub>2</sub>. Observations of photodissociation regions have shown that the emission from aromatic band carriers and from fluorescently-excited H<sub>2</sub> just outside photoionized regions are nearly cospatial, with H<sub>2</sub> sometimes seen to extend slightly deeper into molecular clouds (Sellgren et al. 1990; Graham et al. 1993; Tielens et al. 1993; Brooks et al. 2000; Habart et al. 2003). Cospatiality might be expected since both species can be excited by FUV photons. Aromatic band carriers can also be excited by lower-energy photons in the ultraviolet and optical, but with smaller absorption cross-sections (see Li & Draine 2001), so that FUV photons will dominate the excitation whenever massive stars are present. H<sub>2</sub> is however dissociated by FUV photons between 11.3 and 13.6 eV where it is not self-shielded, whereas PAHs survive the absorption of these photons. Therefore, in the case of relatively dense PDRs (associated with molecular clouds), where collisional heating is expected to be the major origin of the H<sub>2</sub> rotational lines, H<sub>2</sub> emission should peak at slightly higher optical depth than aromatic bands, in the transition layer between atomic and molecular hydrogen,

with  $A_V > 1$ . In addition, PAHs probably cannot be excited as deep into molecular clouds as  $\text{H}_2$ , because at sufficiently high densities they will be coagulated onto grain mantles on short timescales (Boulanger et al. 1990). If photodissociation regions dominate the excitation of  $\text{H}_2$ , as consistent with the above results, a tight relation between aromatic band emission and rotational  $\text{H}_2$  emission can arise only if the physical conditions in PDRs, especially the  $G_0/n$  ratio, are relatively uniform, because  $\text{H}_2$  fluxes and PAH fluxes depend in very different ways on these two parameters. The condition of relatively constant  $G_0/n$  ratios seems verified in the present sample at least for the average emission within kiloparsec-scale regions (see above). Based on the modelling of [CII] and [OI] emission, Malhotra et al. (2001) proposed that a regulation of  $G_0/n$  might be achieved at the scale of individual PDRs by expanding HII regions in pressure equilibrium with their surrounding PDRs.

A correlation was previously claimed by Mouri et al. (1990) between the  $3.3\ \mu\text{m}$  band and the  $v=1-0$  S(1) line at  $2.12\ \mu\text{m}$  for a small sample of starburst and Seyfert galaxies. The dominant source that they propose for  $\text{H}_2$  excitation, following Moorwood & Oliva (1988), is however not photodissociation regions, but shocks in supernova remnants. Using the shock models of Kaufman & Neufeld (1996) to estimate the sum of the S(0) to S(2) transitions (up to 6% of the mechanical power, assuming that its totality is dissipated in molecular clouds), and the population synthesis model of Leitherer et al. (1999) to estimate both the total mechanical power and the FUV luminosity from continuous star formation with a Salpeter initial mass function, shocks alone are in principle able to produce a significant fraction of the observed  $\text{H}_2$  emission, but only if the efficiency of conversion of mechanical power into  $\text{H}_2$  emission is unrealistically high. The rotational line ratios are also inconsistent with shock models, which predict higher temperatures ( $T > 1000\ \text{K}$ ) except for very low shock velocities (in which case the power fraction radiated away by rotational  $\text{H}_2$  lines is lower). If the collective rotational line emission from shocks in supernova remnants is similar to that observed in individual objects such as 3C 391 and IC 443 (Reach et al. 2002), then this mechanism can provide only a modest fraction of the total  $\text{H}_2$  emission. In addition, if  $\text{H}_2$  emission came predominantly from supernova remnants whereas aromatic bands arise mostly in photodissociation regions, the partial disconnection between the two, both temporal and spatial, would manifest itself by a large scatter in the observed relation between  $\text{H}_2$  and PAH fluxes for galaxies with diverse star formation histories, which is not observed.

More recently, Rigopoulou et al. (2002) proposed a relation similar to that presented by Mouri et al. (1990), between the  $7.7\ \mu\text{m}$  aromatic band and the rotational S(1) line in starburst galaxies. Figure 8c not only confirms this result for lower-luminosity galaxies, but also shows that the dispersion for the whole sample of star-forming nuclei is very small, and much smaller with the aromatic bands than with the  $24\ \mu\text{m}$  emission (Fig. 8d), which is dominated by the continuum from transiently-heated very small grains, as well as from

big grains in intense radiation fields. The quantification of the average H<sub>2</sub> to dust power ratios and their dispersions is given in the caption of Fig. 8. The energy coupling between aromatic band carriers and H<sub>2</sub> strongly suggests that both are excited predominantly in photodissociation regions, although they may not come from the exact same layers (at the same optical depths within the clouds). We present further analysis in the next section.

A similar correlation, with a similarly small dispersion, was observed between the [CII] line and aromatic band emission (Helou et al. 2001). This relation suggests that aromatic band carriers are the source of a major part of gas heating in photodissociation regions, via the photoelectric effect, at least at modest radiation field intensities, since [CII] emission is the dominant cooling channel in this case. In the narrow range of physical conditions that seem to apply if the emission from HII nuclei is interpreted in the framework of photodissociation region models (a dynamic range in  $G_0/n$  of only a factor 10), then the same link between aromatic band carriers and H<sub>2</sub> would follow if H<sub>2</sub> were heated in relatively dense photodissociation regions by the PAHs. Our results, however, suggest that in nearly half the star-forming targets, the dominant excitation mechanism of the rotational levels may be fluorescence in low-density regions, so that ortho-para thermalization is not achieved by collisions (see Sect. 6). If the lines are fluorescently excited, the cause underlying the tight relation between H<sub>2</sub> and aromatic band emission may be that both are proportional to the incident far-ultraviolet flux which excited them.

#### 5.4. Cirrus clouds versus PDRs with high radiation field intensities

The tight association between H<sub>2</sub> emission and aromatic bands (Fig. 8c) may be surprising if one assumes that a significant fraction of aromatic band emission arises from diffuse, mostly atomic regions with low radiation field intensities. The infrared emission of such clouds is often termed cirrus (Low et al. 1984; Terebey & Fich 1986). If this were the case, then the scaling of PAH flux with H<sub>2</sub> flux could be explained only if a constant fraction of the total FUV flux escaped PDRs and were absorbed in the more diffuse interstellar medium. We stress that we adopt here the definition of photodissociation regions stated by Hollenbach & Tielens (1997): these are not restricted to the interfaces between bright HII regions and dense molecular clouds, but apply more generally to all the neutral interstellar medium illuminated by FUV photons (with energies between 6 and 13.6 eV). Figure 9 demonstrates the great difficulty of making the idea of an important contribution from the cirrus medium consistent with the data. It shows the flux ratios of H<sub>2</sub> to PAH emission on the one hand, and 24 μm to PAH emission on the other hand, as a function of  $P_{24} \sim \nu_{24}F_{24}/(\nu_{71}F_{71} + \nu_{156}F_{156})$ , estimated within the spectroscopic apertures. The quan-

tity  $P_{24}$  is closely related to  $f_{U>100} = f(L_{\text{dust}} ; U_{\text{rad}} > 100)$ , derived from the modelling of the global spectral energy distributions by Draine et al. (2007), which is the fraction of the total dust luminosity emitted by regions with radiation field intensities  $U_{\text{rad}}$  higher than 100 times the local average value. The dust luminosity fraction of cirrus clouds,  $\sim f_{U<10}$ , can be evaluated as  $1 - c f_{U>100}$ , if one assumes that  $f_{U>100}$  and  $f_{U>10}$  are in constant proportion to each other. Fig. 9 shows that while the  $F_{24}/\text{PAH}$  ratio rises by about one order of magnitude, the  $\text{H}_2/\text{PAH}$  ratio is invariant as a function of  $P_{24}$  or  $f_{U>100}$ . The results obtained by replacing  $P_{24}$  with  $F_{71}/F_{156}$ , which has the same physical significance as  $F_{60}/F_{100}$ , a more traditional indicator of the relative importance of HII regions and cirrus clouds (Helou 1986), are identical.

We conclude that aromatic bands are mostly associated with photodissociation regions (illuminated by FUV photons able to provide  $\text{H}_2$  excitation). In addition, since PAHs are excited not only by FUV photons but also by low-energy photons, the observed constancy of the  $\text{H}_2$  to PAH ratio imposes some restrictions on possible variations of the radiation field hardness. Assuming that cirrus clouds, i.e. PDRs with low radiation field intensities, receive appreciably softer radiation than PDRs with high radiation field intensities, it would be difficult to understand how both types of regions could produce similar  $\text{H}_2/\text{PAH}$  ratios. As a corollary, the hypothesis that cirrus clouds could make a large contribution to  $\text{H}_2$  and PAH emission in our targets, although not definitely ruled out, is not favored. Note that the situation may be different in more quiescent parts of galaxies, not probed by the present sample, and deserves further investigation. Measurement of  $\text{H}_2$  line fluxes in quiescent regions is however challenging, because they depend steeply on the  $G_0$  and  $n$  parameters.

The above does not preclude a large portion of the  $\text{H}_2$  and PAH emission to originate in relatively diffuse molecular gas. Estimates of the optical depth of the  $^{12}\text{CO}(1-0)$  line, over large areas of the Galaxy, indicate that the total molecular medium comprises a substantial diffuse component (Polk et al. 1988). We will see in the next section that our data, for a portion of the targets, do support an important contribution from low-density PDRs to the total warm  $\text{H}_2$  emission.

## 6. Excitation mechanisms in star-forming regions

In sources with purely stellar activity,  $\text{H}_2$  emission is expected to arise in varying proportions from two main energy sources: the far-ultraviolet radiation of OB stars illuminating PDRs; and shocks in supernova remnants or other sources, providing collisional heating. In the first case, the excitation can be thermal, by collisions with gas heated by photoelectrons, or by inelastic collisions with  $\text{H}_2$  molecules pumped by FUV photons. The excitation mech-



anism can also be fluorescence, followed by radiative cascade to the ground vibrational state. Heating by supernova remnants is unlikely to be dominant for two reasons, as we have seen in Sect. 5.3. First, at the low observed temperatures dominating the warm H<sub>2</sub> mass, heating would have to be provided by slow shocks, which are not efficient enough to compete with PDR excitation. Second, variations in star formation histories within the sample, which are shown by Moustakas et al. (2007, in prep.) to be very large (from population synthesis fitting to optical spectra), would produce more scatter than observed in the H<sub>2</sub>/PAH ratio. We conclude that H<sub>2</sub> is heated predominantly by PDRs, and this interpretation is supported both by energetics arguments and by the close association with aromatic band emission (Sect. 5). In this section, we focus on additional constraints on the physical conditions and excitation mechanisms in PDRs (thermal or fluorescent) from the line ratios and excitation diagrams.

### 6.1. Constraints from the temperature distribution

We have seen that the H<sub>2</sub> to far-infrared ratios are consistent with values of  $G_0/n$ , the ratio of the average radiation field intensity to the hydrogen density in PDRs, between about 0.1 and 1. In principle, separate constraints on  $G_0$  and on  $n$  can be obtained from the temperature distribution reflected in the H<sub>2</sub> line ratios (Kaufman et al. 2006). Given the complexity of the surveyed regions, however, the interpretation of the line ratios by comparison with models of a single PDR is severely limited. First, the emission from many distinct PDRs, presumably showing a large range of physical conditions, is averaged within the beam. Second, even though the total emission in the sum of the S(0) to S(3) lines is probably dominated by PDRs, shocks also related to the star formation activity must be present, driven in particular by supernova remnants, protostellar outflows, or turbulence dissipation (Falgarone et al. 2005). Since these shocks are characterized by higher rotational temperatures than PDRs, they would contribute mostly to the S(3) line (Kaufman & Neufeld 1996), and in view of the observed line ratios, negligibly to the lower-lying transitions.

In both cases, the superposition of PDRs of different conditions or of shocks induces a spread in temperatures, and as a consequence these are not reproduced by single PDR models. We compared the line ratios of the star-forming targets with  $OPR_{\text{high T}} = 3$  to the predictions of Kaufman et al. (2006). Although the ranges in  $G_0$  and  $n$  derived from S(2)/S(0) have a broad overlap with those derived from S(1)/S(0), the S(3)/S(1) ratios are inconsistent, indicating more intense radiation fields and higher densities. The least biased tracer of physical conditions in the bulk of the PDRs is thus the S(1)/S(0) ratio (for galaxies with  $OPR_{\text{high T}} = 3$ ). It suggests that the average  $G_0$  varies between about 100 and 5000

(in units of  $1.6 \times 10^{-3} \text{ ergs}^{-1} \text{ cm}^{-2}$ ) and the average  $n$  between about 500 and  $10^4 \text{ cm}^{-3}$ .

## 6.2. Fluorescent excitation

It is surprising that in several objects, the S(0) to S(3) transitions are not thermalized, as indicated by deviations from an apparent  $OPR_{\text{high T}}$  of 3.  $OPR$  values that are different from the equilibrium value at the temperature of the  $\text{H}_2$  gas arise naturally from fluorescent excitation, if the gas density is lower than the critical density for ortho-para equilibration by collisions with H and  $\text{H}^+$  (Sternberg & Neufeld 1999). This is because the ultraviolet absorption lines have a greater optical depth in the ortho states than in the para states, so that the ortho states at a given depth are pumped less. The apparent  $OPR_{\text{high T}}$  values only apply to the states excited by FUV pumping, and do not imply that the true  $OPR_{\text{high T}}$  value (for the J=1 and J=0 states, where most of the gas resides) be different from 3. For a total  $OPR_{\text{high T}}$  of 3, the  $OPR_{\text{high T}}$  of the excited levels is predicted to be close to  $\sqrt{3} = 1.7$ . A thorough review of the phenomenon of selective excitation and its implications for the interpretation of excitation diagrams was provided by Sternberg & Neufeld (1999). Note that radiative decay in the electronic ground state and most collisional (de)excitations always occur at fixed spin number and thus preserve the ortho or para state. Conversion from one state to the other can be accomplished by  $\text{H}_2$  dissociation followed by reformation on dust grains, or by reactions with protons and hydrogen atoms in the gas phase.

Deviations from thermalization had previously remained unseen for rotational lines of extragalactic sources (Rigopoulou et al. 2002; Higdon et al. 2006). In the frame of PDR excitation, values  $OPR_{\text{high T}} < 3$  may be interpreted as arising in FUV-pumped gas with sufficiently low densities to prevent ortho-para equilibration by collisions. However, because this process depends on unknown timescales for gas heating, cooling, dissociation and reformation, we are unable to quantify in a simple way the implied density conditions. Alternatively, it is conceivable that the emission comes from initially cold gas, that has been heated by slow shocks and has not had time to reach the equilibrium value of  $OPR$  (Timmermann 1998; Wilgenbus et al. 2000). For HII nuclei, however, we have seen that photodissociation region excitation is more likely than shock heating (Sect. 5).

Another possibility is that in a fraction of the PDRs within the beam, the photodissociation front is advancing into cold gas (in LTE at time  $t_1$  with  $OPR(t_1) \ll 3$  and  $OPR_{\text{high T}}(t_1) = 3$ ), and the recently-heated gas has not yet had enough time to reach LTE at time  $t_2$  (with  $OPR(t_2) \sim OPR(t_1)$  and  $OPR_{\text{high T}}(t_2) < 3$  according to Equ. 2). This is the interpretation favored by Fuente et al. (1999) and Habart et al. (2003) to explain the non-LTE ortho to para ratios observed in the PDRs of NGC 7023 and  $\rho$  Ophiuchi,

respectively. In this scenario, the observed portion of the interstellar medium would have to contain a much larger fraction of non-equilibrium PDRs in targets with low  $OPR_{\text{high T}}$  than in targets with  $OPR_{\text{high T}} = 3$ , and the underlying reason would be unclear. It is also unknown whether the timescales involved in ortho-para equilibration in our sources are long enough for this scenario to be viable.

Figure 10 shows the derived  $OPR_{\text{high T}}$  values (fixed to 3 whenever the temperatures derived from each pair of adjacent transitions were compatible with this assumption) as a function of the total brightness of the S(0) to S(2) lines. We have here included LINERs and Seyferts because they do not display a different behavior in this diagram. Under the hypothesis of photodissociation region excitation, and assuming first that sites of star formation occupy a constant fraction of the observing beam, sources with low  $\text{H}_2$  brightnesses should consist of regions with both low densities and low radiation field intensities, while sources with the highest  $\text{H}_2$  brightnesses should include a greater fraction of high-density, high-radiation regions. In this simplified view, low values of  $OPR_{\text{high T}}$ , indicating that  $\text{H}_2$  is not thermalized by collisions, could be obtained only in the low-brightness sources, as seen generally in Fig. 10. Variations in the beam filling factor by sites of star formation would then induce a horizontal scatter, which is indeed very large. Fig. 10 also shows  $OPR_{\text{high T}}$  as a function of the average surface brightness in the  $24\ \mu\text{m}$  band. The latter quantity incorporates a significant contribution from HII regions, and should be dominated by variations in radiation field intensity, rather than variations in gas density. The fact that the horizontal spread is larger in this diagram than in the diagram involving the  $\text{H}_2$  brightness supports our tentative interpretation in terms of density effects. In addition,  $OPR_{\text{high T}}$  does not show any variation as a function of  $F_{71}/F_{156}$ , and only a weak tendency to increase with the quantity  $P_{24}$  discussed in the previous section. Both these quantities are indicators of the average radiation field intensity; they are observed to be generally correlated with the gas density, but only weakly. Although the data do not allow us to truly estimate average densities in photodissociation regions, they suggest that in a substantial number of the observed nuclear regions, the emission can be dominated by low density gas, relative to well-studied Galactic PDRs and starburst galaxies. New modelling is required to quantify the conditions under which rotational lines indicate  $OPR_{\text{high T}} < 3$  in PDRs.

The  $\text{H}_2$  line ratios, compared with the PDR model of Kaufman et al. (2006), do not indicate lower gas densities, on average, in galaxies with  $OPR_{\text{high T}} < 3$  than in galaxies with  $OPR_{\text{high T}} = 3$ . However, the densities estimated in this way are averages within the whole beam, and the densest and warmest regions have a greater weight, because they are more luminous in  $\text{H}_2$ ; if a large spread in densities exists, with both dense clumps and diffuse PDRs, an increase in the diffuse fraction may not be easily detectable in the average density, while still leaving an imprint on  $OPR_{\text{high T}}$ . We note that the proton density should also

play an important role, since ortho-para conversion is effected by collisions with H and H<sup>+</sup>. Whether galaxies with  $OPR_{\text{high T}} < 3$  have lower proton densities in molecular clouds, due to reduced ionization by cosmic rays, is in principle testable with radio continuum observations of synchrotron radiation from cosmic ray electrons. The number of sample galaxies with adequate data, at sufficiently high angular resolution, is however too small to apply this test.

We have seen in Section 3.3 that extinction effects are unlikely to modify our results in a statistical sense. The fitted  $OPR_{\text{high T}}$  values are correlated with the brightness of the dust emission, in the aromatic bands, the 24  $\mu\text{m}$  continuum and total infrared emission (not shown here but similar to Fig. 10), which is a further indirect argument against extinction effects being responsible for low  $OPR_{\text{high T}}$  values. If extinction played a significant role, then low  $OPR_{\text{high T}}$  values would be seen preferentially in bright and compact regions, which is not the case.

The galaxies for which we derive the lowest  $OPR_{\text{high T}}$  values (and among the lowest H<sub>2</sub> surface brightnesses, regardless of the constraint on the  $T_1$  temperature) are NGC 337, NGC 1705 ( $OPR_{\text{high T}}$  is a lower limit), NGC 2915, NGC 4552 ( $OPR_{\text{high T}}$  is a lower limit) and NGC 7793. The first three are dwarf galaxies, NGC 7793 a very late-type spiral and NGC 4552 is a small elliptical galaxy classified as LINER, with the smallest infrared brightness of the whole sample. Since smaller gas densities are expected in general in dwarf galaxies than in the central regions of massive galaxies, except in blue compact dwarfs such as Mrk 33, this finding is consistent with our interpretation of small  $OPR_{\text{high T}}$  values in terms of low density. We may also remark that in NGC 1705, which is a starburst, no ultracompact HII region was detected in the radio (Johnson et al. 2003); thus no conflict exists with the hypothesis that the photodissociation regions in NGC 1705 have low densities.

## 7. Excitation mechanisms in LINER and Seyfert nuclei

A large number of the galaxies classified as LINERs or Seyferts deviate significantly from the relations discussed in Section 5, in having a strong excess of H<sub>2</sub> emission with respect to all the other tracers used here, not only aromatic bands but also [SiII], the 24  $\mu\text{m}$  flux and the total infrared emission (Fig. 8), arguing for an alternative excitation mechanism in these galaxies. The average and dispersion of each power ratio are given separately for HII nuclei and for LINER and Sy nuclei in the figure caption. Given these results, the quantities by which the two categories are most clearly separated are the H<sub>2</sub> to aromatic band ratio and the H<sub>2</sub> to [SiII] ratio. In particular, nuclear regions with  $F(S0 - S2)/F_{7.9\text{dust}} > 10^{-1.94}$  are likely to be of the LINER or Sy type at the 99% confidence level.

We thus define an excess of H<sub>2</sub> emission, with respect to HII nuclei, based on the observed relation with aromatic bands. We choose as the maximal H<sub>2</sub> power associated purely with star formation the quantity  $10^{-1.94} \times L_{7.9 \text{ dust}}$ . Aromatic band carriers are thought to be destroyed in intense radiation field environments, such as HII region cores (Giard et al. 1994) and ionized regions around Seyfert nuclei (Désert & Dennefeld 1988; Voit 1992). They would however survive where H<sub>2</sub> is not dissociated, so that an enhancement of the H<sub>2</sub>/PAH ratio is more likely caused by a genuine excess of H<sub>2</sub> emission, in conditions where PAHs are not excited, rather than normal H<sub>2</sub> emission occurring where PAHs would have been destroyed. Furthermore, H<sub>2</sub> emission is seen in excess not only with respect to aromatic bands, but also with respect to [SiII] and other dust tracers (24 μm and total infrared emission). [SiII] emission may be depressed because the ionization state of silicon becomes higher, but the dust continuum cannot be suppressed like aromatic bands. Our empirical quantification of the H<sub>2</sub> excess is intended to extract the part of H<sub>2</sub> emission that cannot originate in photodissociation regions. The excitation mechanism that is required to account for this excess, while not exciting PAHs, is either X-ray irradiation or shock heating.

We discuss LINER and Sy nuclei as a single category, because no detectable difference exists in their H<sub>2</sub> properties. This may be expected for several reasons: small-number statistics; the fact that the classification of low-luminosity AGNs can be ambiguous as it depends on the aperture size in particular; and the fact that the source of H<sub>2</sub> excitation might not be directly linked to the nuclear activity, as the results discussed below suggest.

### 7.1. Heating by X-rays from an active galactic nucleus

The idea that a nuclear X-ray source may modify the chemistry and excitation of the surrounding molecular clouds through sufficiently large column densities as to offer a convenient way to identify active nuclei hidden by dust has recently received much attention. Models predict, in particular, unusual ratios of tracers of dense molecular gas (Lepp & Dalgarno 1996; Meijerink & Spaans 2005), consistent with observations of NGC 1068 (Usero et al. 2004). X-ray excitation would also manifest itself in the properties of H<sub>2</sub>, the most abundant molecule. To test the hypothesis that the additional excitation in galaxies showing a significant excess of H<sub>2</sub> emission with respect to aromatic bands (Figure 8c) is predominantly produced by nuclear X-rays, we have compiled estimated X-ray fluxes in the 2-10 keV band obtained from Chandra observations. The data and references are summarized in Table 6. The H<sub>2</sub> excess is shown as a function of the X-ray luminosity of the nucleus in Fig. 11. Here, galaxies with no H<sub>2</sub> excess according to our definition are shown below the dashed line.

The spread in Figure 11 is very large. Since the available X-ray measurements do

not isolate the hard X-ray component, but include soft emission, a substantial part can be thermal emission from supernova remnants, as opposed to power-law emission from the activity related to a central supermassive black hole. Many HII nuclei with no H<sub>2</sub> excess indeed have X-ray luminosities that are comparable to those of LINER and Seyfert nuclei. On the other hand, some galaxies, in particular NGC 3627, NGC 4569 and NGC 5866, are very luminous in H<sub>2</sub> but have only modest X-ray to H<sub>2</sub> luminosity ratios compared with other H<sub>2</sub>-excess galaxies. We shall assume a power-law spectrum with a standard photon index of  $-1.8$ , as adopted by Ho et al. (2001), to extrapolate the total X-ray luminosity to lower photon energies. Under this assumption, the luminosity between 0.2 and 10 keV is only two times the luminosity between 2 and 10 keV compiled here. Up to 10% of the intrinsic 1-10 keV luminosity may emerge in the sum of all H<sub>2</sub> transitions (Lepp & McCray 1983), whereas in our comparison we have summed only the three rotational lines S(0) to S(2). Therefore, pure X-ray excitation of the excess H<sub>2</sub> emission should not be energetically possible even in objects such as NGC 4579 and NGC 5195, which have relatively low H<sub>2</sub> to X-ray luminosity ratios in Fig. 11. Even if the intrinsic emission in soft X-rays were underestimated by a large factor, X-ray excitation would still be unlikely to dominate in most cases. The apparent trend of increasing excess H<sub>2</sub> luminosities with increasing X-ray luminosities in Fig. 11 does not imply a direct excitation of H<sub>2</sub> by X-rays, since it could also be understood in the frame of multi-phase shocks, produced by supernova remnants or by starburst winds (see below). From near-infrared line ratios, Davies et al. (2005) also reached the conclusion that X-rays do not contribute significantly to the excitation of H<sub>2</sub> in a small sample of active galactic nuclei, within regions smaller than those sampled here (about 100 pc).

Supernova remnants are another important source of X-rays. More than half their mechanical energy can be converted to X-rays, mostly with energies below  $\sim 500$  eV (Draine & Woods 1991), but the X-ray power is converted to H<sub>2</sub> emission in the S(0) to S(2) lines with low efficiencies, less than  $10^{-3}$  (Draine & Woods 1990). Therefore, in view of the estimates presented in Fig. 12 (see the next Section for details), heating of H<sub>2</sub> by X-rays from supernova remnants is completely negligible.

## 7.2. Heating by shocks

For galaxies such as NGC 3627 and NGC 4569, with very high H<sub>2</sub> to X-ray luminosity ratios, an efficient mechanism has to be invoked to account for the H<sub>2</sub> brightness. It has been shown recently that galactic shocks can convert a very large fraction of the kinetic energy into rotational H<sub>2</sub> emission without producing a lot of X-rays (Appleton et al. 2006). We thus propose that large-scale shocks play a major role. NGC 3627 is an interacting

galaxy in the Leo Triplet, characterized by severe morphological and kinematical distortions and, in the center, a massive molecular gas concentration in the form of an asymmetric bar-like structure (Zhang et al. 1993), with a peak in the stellar velocity dispersion that is shifted by about  $3''$  from the nucleus, and probably strong gas inflow into the nucleus (Afanasiev & Sil’chenko 2005). NGC 4569 is a starburst LINER (Maoz et al. 1998) having recently produced a large number of supernova explosions, triggering the expansion of X-ray,  $H\alpha$  (Tschöke et al. 2001) and synchrotron emission lobes (Chyzy et al. 2006), and also has a circumnuclear ring of molecular gas with strong non-circular motions (Nakanishi et al. 2005). The situation could be similar to NGC 4945, where  $H_2$  emission is seen to follow the innermost part of a starburst outflow, with an extent of  $\sim 200$  pc (Moorwood et al. 1996). These two examples thus support the idea of a dominant excitation by shocks, triggered by dynamical perturbations and by a starburst wind, respectively.

In order to test the possibility that the excess  $H_2$  emission be caused by supernova remnant shocks, we estimated the required heating efficiencies in the following way. We first computed the star formation rate required to account for the excess  $H_2$  emission, assuming constant star formation, if all the mechanical power is converted to the power emitted by the sum of the S(0) to S(2) lines (using the Starburst99 population synthesis model of Leitherer et al. 1999). We then compared this with the star formation rate estimated from the  $24\ \mu\text{m}$  luminosity, using the calibration of Wu et al. (2005) for the same initial mass function. The ratio of the two rates gives an order-of-magnitude estimate of the  $H_2$  heating efficiency  $\eta$  that is needed if supernova shocks are invoked as the dominant heating mechanism, and is shown in Fig. 12 as a function of the excess  $H_2$  luminosity. In this simplified computation, the  $H_2$  excess in the star-forming regions B and C of NGC 6822 can easily be attributed to supernova remnant shocks, as they would imply efficiencies of at most 3%. Besides these regions, supernova remnants are also a sufficient heating source of the excess  $H_2$  for at least six of the LINER and Sy nuclei. In NGC 5195, in particular, the necessary heating efficiency (including the fraction of the total mechanical power injected and absorbed in molecular clouds) is very small. Galaxies such as NGC 4450 and NGC 4579, on the other hand, unambiguously require much more power than available in supernova remnants. The excitation source could be shocks triggered by cloud collisions induced by gravitational perturbations, maybe in combination with X-rays.

NGC 1316 (Fornax A) stands out in Fig. 7 as having the highest  $T_1$  temperature ( $T_1 > 280$  K) and the lowest mass fraction in warm phase ( $< 1\%$ ), and the lowest levels up to J=4 are characterized by a single temperature. In this particular case,  $H_2$  may be heated by fast shocks caused by the powerful jet that has produced large-scale radio lobes, and that is observed in the form of knots in the inner kiloparsec, which may be a signature of interaction with the interstellar medium, i.e. shocks (Geldzahler & Fomalont 1984). Excitation by far-

ultraviolet radiation and X-rays or by slow shocks would produce a large quantity of  $\text{H}_2$  at temperatures between 100 and 300 K (Burton et al. 1992; Maloney et al. 1996), which is ruled out by the data. NGC 1316 is the only galaxy of the sample for which the excess  $\text{H}_2$  may be heated by an AGN jet.

Of the 16  $\text{H}_2$ -excess galaxies, all have excitation diagrams consistent with  $OPR_{\text{high T}} = 3$  (at least in the free- $T_1$  fits), except NGC 1266, NGC 5866 and NGC 4125. NGC 1266 may be unusual simply by virtue of its significant optical depth, as discussed in Section 3.3. This predominance of  $OPR_{\text{high T}} = 3$ , combined with more elevated temperatures than in HII nuclei on average (Fig. 7) would be consistent with the excess  $\text{H}_2$  emission originating in shocks where ortho-para equilibration is fast. It remains however difficult to test this idea, and to identify the source of these shocks.

## 8. Summary and conclusions

We present the measurements and results of a survey of the four lowest-energy rotational transitions of  $\text{H}_2$ , S(0) to S(3), in a local sample of 57 galaxies, from the SINGS program. For three galaxies in this sample, higher-energy transitions, up to S(7), could be measured. Characterizing the amount and physical conditions of the warm molecular hydrogen phase traced by these lines is of prime interest, because molecular hydrogen represents a major mass fraction of the interstellar medium of normal galaxies, and the warm phase itself (gas heated to temperatures of  $\approx 100$  to 1000 K) can constitute a substantial fraction of the total  $\text{H}_2$ .

The emission is measured over areas of median size 0.9 kpc, thus including a large number of distinct star formation sites and molecular clouds. The sample comprises mostly nuclear regions (in 47 massive galaxies and 9 dwarf galaxies), of which 45% are optically classified as LINER or Seyfert, as well as 10 extranuclear star-formation complexes within a dwarf galaxy and a spiral. With respect to earlier studies of molecular hydrogen emission in galaxies, and particularly rotational lines, which had focussed on very bright systems (nearby starbursts and AGNs, as well as ultraluminous galaxies), this paper provides results on the average properties of warm  $\text{H}_2$  of relatively faint systems, more representative of the general population of galaxies.

Perhaps the most significant observational results (detailed below) are: (1) the tight correlation of the powers emitted by the sum of the S(0) to S(2) lines and by aromatic bands, and the fact that the  $F(\text{S0-S2})/\text{PAH}$  ratio is insensitive to the marked variations in average radiation field intensities existing in our sample; (2) the existence of non-equilibrium ortho



to para ratios in the rotational levels, that are weakly correlated with the surface brightness of the H<sub>2</sub> lines. These results call for further modelling in order to be better understood.

**Masses and column densities:** The total masses of warm H<sub>2</sub> within our apertures range between  $10^5 M_{\odot}$  and close to  $3 \times 10^8 M_{\odot}$  in galaxy nuclear regions. In star formation complexes of nearby dwarf galaxies (NGC 2915 and NGC 6822), we probe warm H<sub>2</sub> masses down to a few  $10^3 M_{\odot}$  within equivalent diameters of 60 to 250 pc. The mass densities range between 0.2 and  $30 M_{\odot} \text{ pc}^{-2}$ . The column densities that we derive are on average of the same order of magnitude as the column densities observed in individual Galactic photodissociation regions. For systems in which H<sub>2</sub> is predominantly excited by PDRs, assuming they have similar characteristics to the Orion Bar, this implies that they fill most of the observing beam.

**H<sub>2</sub> mass fraction in the warm phase ( $T \geq 100 \text{ K}$ ):** Under a conservative assumption about the distribution of temperatures, we find that the warm H<sub>2</sub> gas makes up between 1% and more than 30% of the total H<sub>2</sub>. For star-forming galaxies, the median mass fraction in the warm phase is 10%. The column density nevertheless has a steep inverse dependence on the temperature, and we cannot rule out that the unconstrained cool H<sub>2</sub> component ( $70 \text{ K} < T < 100 \text{ K}$ ) might in some cases dominate the S(0) and S(1) emission and account for most of the H<sub>2</sub> mass.

**PDR excitation in star-forming regions:** In HII nuclei, we observe a remarkably narrow range of H<sub>2</sub> to aromatic band flux ratio. This result argues for photodissociation regions providing most of the power used for H<sub>2</sub> excitation, since aromatic bands are known to arise predominantly from these regions, as defined by Hollenbach & Tielens (1997), that include all the neutral interstellar medium illuminated by FUV photons. Two main excitation mechanisms can be at work simultaneously in PDRs: pumping by FUV photons, followed by fluorescent decay to the ground electronic state, and collisions with hydrogen atoms and molecules heated by photoelectrons or FUV pumping. Comparison with the predictions of PDR models for the ratio of H<sub>2</sub> to FUV luminosity indicates a narrow range of average physical conditions,  $G_0/n$  between 0.1 and 1 with the radiation field intensity  $G_0$  in units of  $1.6 \times 10^{-3} \text{ erg s}^{-1} \text{ cm}^{-2}$  and the hydrogen nucleus density  $n$  in units of  $\text{cm}^{-3}$ . The sum of the S(0) to S(2) transitions represents between  $2.5 \times 10^{-4}$  and  $7.5 \times 10^{-4}$  of the total infrared power, and on average 30% of the [SiII] line power. The observed temperatures suggest that the average  $G_0$  varies between about 100 and 5000, and the average  $n$  between about 500 and  $10^4 \text{ cm}^{-3}$ . This seems to imply that H<sub>2</sub> rotational line emission comes mostly from molecular clouds illuminated by OB associations. We have seen however that the estimator of these parameters is biased (see Section 6.1), and a non-negligible contribution from less dense regions with less intense radiation fields is not excluded.

**Evidence for fluorescence in star-forming regions:** Previous surveys of rotational lines in galaxies had not revealed any departure from thermalization of  $\text{H}_2$ , which is a consequence of the relatively low critical densities of the lower rotational transitions. By contrast, we find that nearly half the targets in our sample deviate significantly from local thermodynamic equilibrium in having apparent ortho to para ratios  $OPR_{\text{high T}}$  lower than the equilibrium value of three. We have seen that this result cannot be an artifact caused by extinction effects. Low values of  $OPR_{\text{high T}}$  may thus be interpreted as evidence of fluorescent excitation, which naturally leads to low ortho to para ratios in the excited states, occurring in regions of sufficiently low density that ortho-para equilibration by collisions is incomplete. The fraction of relatively diffuse molecular gas in normal galaxies could thus be far from negligible. In order to test this idea, it would be desirable to obtain independent estimates of the average gas density in PDRs. In the absence of any robust estimate of this quantity, we used the surface brightness of the sum of the S(0) to S(2) transitions as a tracer of the coupled variations of  $n$  and  $G_0$  to show that the data are compatible with an interpretation of low apparent  $OPR_{\text{high T}}$  values in terms of low density PDRs. In particular, the lowest values occur preferentially in very late type galaxies, that may have a more diffuse interstellar medium than earlier-type galaxies. In the present sample, close to half the star-forming targets have low  $OPR_{\text{high T}}$  values. We infer that fluorescence can be the predominant excitation mechanism of rotational  $\text{H}_2$  lines in normal star-forming galaxies. In more active galaxies, however, collisional excitation in photodissociation regions is likely to overtake fluorescence, even though the latter still contributes to gas heating. Alternatively, low  $OPR_{\text{high T}}$  values may be caused by non-equilibrium PDRs, in which initially cold gas, recently reached by the photodissociation front and heated, has not had enough time to adjust its ortho to para ratio. A disadvantage of this scenario is that it does not explain why PDRs would have systematically different properties in targets with  $OPR_{\text{high T}} < 3$  and in those with  $OPR_{\text{high T}} = 3$ . We thus emphasize that the cause of non-equilibrium ortho to para ratios in the rotational levels is not well understood at present.

**Differences between HII and LINER/Sy nuclei:** Despite our large observing beam implying that the emission from the immediate vicinity of the nucleus is diluted in the emission from extended areas decoupled from nuclear activity, a large fraction of nuclei classified as LINER or Seyfert distinguish themselves from purely star-forming nuclei in several ways. In a statistical sense, the temperatures of the warm  $\text{H}_2$  phase are slightly higher, and as a corollary the mass fractions of warm to total  $\text{H}_2$  are lower (with a median of 4% instead of 10% for HII nuclei). The correlation between  $\text{H}_2$  and aromatic band emission observed in HII nuclei also breaks down in LINER and Sy nuclei. A large number of them have excess emission in the  $\text{H}_2$  lines, with respect to aromatic bands and to [SiII], which in general is the brightest cooling line in the mid-infrared range (with possible contributions

from both PDRs and HII regions as well as X-ray irradiated gas), and, to a lesser degree, with respect to the total infrared emission. The fact that less contrast between the different nuclear categories is seen in the  $H_2/TIR$  ratio may partly be due to the fact that estimating the  $70\ \mu\text{m}$  and  $160\ \mu\text{m}$  fluxes within our small apertures requires a large extrapolation.

**Threshold for nonstellar excitation:** We propose that nuclear ratios  $F(S0 - S2)/F_{7.9\ \text{dust}} > 10^{-1.94}$ , with observables defined as in our study, are indicative of the LINER and Sy categories. It should however be kept in mind that some sources belonging to these classes are indistinguishable from HII nuclei. This could be thought of as mainly a distance and beam dilution effect; but the  $H_2/PAH$  ratio behaves contrary to this expectation, since in the present sample it shows a correlation with distance, rather than an anti-correlation. None of the quantities derived in this paper shows any dependence on distance, with this exception of the  $H_2$  excess in LINER and Sy nuclei. The fact that it tends to increase with distance could be a selection bias, but it also suggests that the  $H_2$  excess is in general spatially extended. We may also remark that for galaxies like NGC 1377, whose dust emission is interpreted as dominated by an extremely young and opaque starburst (Roussel et al. 2006), the above criterion to select LINER and Sy nuclei will not be applicable.

**Shock excitation in LINER/Sy nuclei:** We interpret the differences in  $H_2/PAH$  ratios as a genuine excess of  $H_2$  emission (and not a deficit of the other tracers while preserving  $H_2$ ), i.e. as requiring at least one additional mechanism to excite  $H_2$  molecules with respect to PDR heating. Excitation by nuclear X-rays seems implausible, as models predict much lower heating efficiencies than what would be necessary to account for the estimated  $H_2$  excess, compared with X-ray luminosities derived from Chandra observations. We thus favor excess heating by large-scale shocks, caused either by the collective effect of supernovæ in an aging starburst, or by dynamical perturbations. An order-of-magnitude estimate suggests that supernova remnant shocks can easily account for the  $H_2$  excess of a fraction of the LINER and Sy nuclei (for example NGC 5195), but do not provide enough mechanical power for the galaxies with the highest  $H_2/PAH$  ratios. For the latter, shocks triggered by dynamical perturbations are the best candidate to supply the excess  $H_2$  heating. In one case, it is conceivable that the excess  $H_2$  emission may be produced by the interaction of a nuclear jet with the interstellar medium, namely in NGC 1316 (Fornax A). This target has the warmest  $H_2$  of the whole sample, with no evidence for a cool ( $T < 300\ \text{K}$ ) component, the lowest mass fraction in the warm phase, and is remarkable for its large-scale radio lobes.

**Ortho-para thermalization in LINER/Sy nuclei:** Consistent with the hypothesis that the additional  $H_2$  may be caused by shocks in which ortho-para equilibration is fast, the excitation diagrams of most galaxies with excess  $H_2$  emission are consistent with ortho-para thermalization ( $OPR_{\text{high T}} = 3$ ), with three exceptions: NGC 1266, where we have

seen that the apparent  $OPR_{\text{high T}} < 3$  may be a result of high optical depth toward H<sub>2</sub>, NGC 5866 and NGC 4125. Note that the latter two galaxies are observed edge-on, so that  $OPR_{\text{high T}} < 3$  could also be an extinction artefact; this would however imply that the optical depth derived from the silicate bands be underestimated by large factors (see Sect. 3.3). If excess H<sub>2</sub> emission originated from shocks with  $OPR_{\text{high T}} < 3$  in the latter two galaxies, it would imply that the gas was initially cold and had not had time to reach equilibrium.

The rotational H<sub>2</sub> lines are most often fainter than the forbidden lines in the same wavelength range, in particular [SiII], [NeII], [NeIII] and [SIII] (with no example of an H<sub>2</sub>-dominated line spectrum in the present sample), but are among the dominant coolants of molecular gas and provide important constraints on the excitation of the warm molecular interstellar medium, where a large mass fraction of the gas resides in normal galaxies. The results presented here are assumed to be representative of moderate-luminosity galaxies of all types, and can serve as a comparison point for future studies of distant galaxies.

Support for this work, part of the Spitzer Space Telescope Legacy Science Program, was provided by NASA through an award issued by the Jet Propulsion Laboratory, California Institute of Technology under NASA contract 1407. We thank Adam Leroy for a useful discussion.

## REFERENCES

- Aalto, S., Booth, R.S., Black, J.H., & Johansson, L.E. 1995, A&A, 300, 369 (A95)
- Afanasiev, V.L., & Sil’chenko, O.K. 2005, A&A, 429, 825
- Albrecht, M., Chini, R., Krügel, E., Müller, S.A., & Lemke, R. 2004, A&A, 414, 141 (A04)
- Allers, K.N., Jaffe, D.T., Lacy, J.H., Draine, B.T., & Richter, M.J. 2005, ApJ, 630, 368
- Appleton, P.N., Xu, K.C., Reach, W., et al. 2006, ApJ, 639, L51
- Bajaja, E., Krause, M., Wielebinski, R., & Dettmar, R.J. 1991, A&A, 241, 411 (B91)
- Bally, J., & Lane, A.P. 1982, ApJ, 257, 612
- Beck, S.C., Lacy, J.H., & Geballe, T.R. 1979, ApJ, 234, L213
- Beckwith, S., Gatley, I., & Persson, S.E. 1978, ApJ, 219, L33
- Black, J.H., & Dalgarno, A. 1976, ApJ, 203, 132

- Böker, T., Lisenfeld, U., & Schinnerer, E. 2003, *A&A*, 406, 87 (B03)
- Boulanger, F., Falgarone, E., Puget, J.L., & Helou, G. 1990, *ApJ*, 364, 136
- Braine, J., Combes, F., Casoli, F., et al. 1993, *A&AS*, 97, 887 (B93)
- Brooks, K.J., Burton, M.G., Rathborne, J.M., Ashley, M.C., & Storey, J.W. 2000, *MNRAS*, 319, 95
- Burton, M.G., Hollenbach, D.J., & Tielens, A.G. 1992, *ApJ*, 399, 563
- Burton, M.G., Brand, P.W., Geballe, T.R., & Webster, A.S. 1989, *MNRAS*, 236, 409
- Chandar, R., Bianchi, L., & Ford, H.C. 2000, *AJ*, 120, 3088
- Chyzy, K.T., Soida, M., Bomans, D.J., et al. 2006, *A&A*, 447, 465
- Claussen, M.J., & Sahai, R. 1992, *AJ*, 103, 1134 (C92)
- Contini, T., Wozniak, H., Considere, S., & Davoust, E. 1997, *A&A*, 324, 41 (C97)
- Curran, S.J., Polatidis, A.G., Aalto, S., & Booth, R.S. 2001, *A&A*, 368, 824 (C01)
- Dale, D.A., Gil de Paz, A., Gordon, K.D., et al. 2007, *ApJ*, 655, 863
- Dale, D.A., Smith, J.D., Armus, L., et al. 2006, *ApJ*, 646, 161
- Dale, D.A., & Helou, G. 2002, *ApJ*, 576, 159
- Dame, T.M., Hartmann, D., & Thaddeus, P. 2001, *ApJ*, 547, 792
- Davies, R.I., Sternberg, A., Lehnert, M.D., & Tacconi-Garman, L.E. 2005, *ApJ*, 633, 105
- Davies, R.I., Sternberg, A., Lehnert, M., & Tacconi-Garman, L.E. 2003, *ApJ*, 597, 907
- Désert, F.X., & Dennefeld, M. 1988, *A&A*, 206, 227
- Draine, B.T., & Li, A. 2007, *ApJ*, 657, 810
- Draine, B.T., Dale, D.A., Bendo, G., et al. 2007, *ApJ*, in press (astro-ph/0703213)
- Draine, B.T., & Woods, D.T. 1991, *ApJ*, 383, 621
- Draine, B.T., & Woods, D.T. 1990, *ApJ*, 363, 464
- Draine, B.T., Roberge, W.G., & Dalgarno, A. 1983, *ApJ*, 264, 485

- Elfhag, T., Booth, R.S., Hoeglund, B., Johansson, L.E., & Sandqvist, A. 1996, *A&AS*, 115, 439 (E96)
- Falgarone, E., Verstraete, L., Pineau des Forêts, G., & Hily-Blant, P. 2005, *A&A*, 433, 997
- Fazio, G., Hora, J.L., Allen, L.E., et al. 2004, *ApJS*, 154, 10
- Fuente, A., Martin-Pintado, J., Rodriguez-Fernandez, N.J., Rodriguez-Franco, A., de Vicente, P., & Kunze, D. 1999, *ApJ*, 518, L45
- Gatley, I., Hasegawa, T., Suzuki, H., et al. 1987, *ApJ*, 318, L73
- Gatley, I., Beattie, D.H., Lee, T.J., Jones, T.J., & Hyland, A.R. 1984, *MNRAS*, 210, 565
- Gautier, T.N., Fink, U., Treffers, R.R., & Larson, H.P. 1976, *ApJ*, 207, L129
- Geldzahler, B.J., & Fomalont, E.B. 1984, *AJ*, 89, 1650
- Giard, M., Bernard, J.P., Lacombe, F., Normand, P., & Rouan, D. 1994, *A&A*, 291, 239
- Gordon, M.A. 1991, *ApJ*, 371, 563 (G91)
- Graham, J.R., Serabyn, E., Herbst, T.M., et al. 1993, *AJ*, 105, 250
- Habart, E., Boulanger, F., Verstraete, L., Pineau des Forêts, G., Falgarone, E., & Abergel, A. 2003, *A&A*, 397, 623
- Helfer, T.T., Thornley, M.D., Regan, M.W., et al. 2003, *ApJS*, 145, 259
- Helfer, T.T., & Blitz, L. 1993, *ApJ*, 419, 86 (H93)
- Helou, G., Roussel, H., Appleton P., et al. 2004, *ApJS*, 154, 253
- Helou, G., Malhotra, S., Hollenbach, D.J., Dale, D.A., & Contursi, A. 2001, *ApJ*, 548, L73
- Helou, G., Khan, I.R., Malek, L., & Boehmer, L. 1988, *ApJS*, 68, 151
- Helou, G. 1986, *ApJ*, 311, L33
- Herbst, T.M., Beckwith, S.V., Glindemann, A., Tacconi-Garman, L.E., Kroker, H., & Krabbe, A. 1996, *AJ*, 111, 2403
- Herbst, T.M., Graham, J.R., Tsutsui, K., Beckwith, S., Matthews, K., & Soifer, B.T. 1990, *AJ*, 99, 1773
- Higdon, S.J., Armus, L., Higdon, J.L., Soifer, B.T., & Spoon, H.W. 2006, *ApJ*, 648, 323

- Ho, L.C., Feigelson, E.D., Townsley, L.K., Sambruna, R.M., Garmire, G.P., Brandt, W.N., et al. 2001, *ApJ*, 549, L51 (HF01)
- Ho, L.C., Filippenko, A.V., & Sargent, W.L. 1997, *ApJS*, 112, 315
- Hollenbach, D.J., & Tielens, A.G. 1997, *ARA&A*, 35, 179
- Horellou, C., Black, J.H., van Gorkom, J.H., Combes, F., van der Hulst, J.M., & Charmandaris, V. 2001, *A&A*, 376, 837 (H01)
- Houck, J.R., Roellig, T.L., van Cleve, J., et al. 2004, *ApJS*, 154, 18
- Hubble, E. 1925, *ApJ*, 62, 409
- Huber, K.P., & Herzberg, G. 1979, *Constants of diatomic molecules*, ed. van Nostrand
- Israel, F.P. 2005, *A&A*, 438, 855 (I05)
- Israel, F.P., Baas, F., Rudy, R.J., Skillman, E.D., & Woodward, C.E. 2003, *A&A*, 397, 87
- Israel, F.P., & Baas, F. 1999, *A&A*, 351, 10 (I99)
- Israel, F.P., Tacconi, L.J., & Baas, F. 1995, *A&A*, 295, 599 (I95)
- Johnson, K.E., Indebetouw, R., & Pisano, D.J. 2003, *AJ*, 126, 101
- Kaufman, M.J., Wolfire, M.G., & Hollenbach, D.J. 2006, *ApJ*, 644, 283
- Kaufman, M.J., & Neufeld, D.A. 1996, *ApJ*, 456, 611
- Kenney, J.D., & Young, J.S. 1988, *ApJS*, 66, 261 (K88)
- Kennicutt, R.C., et al. 2003, *PASP* 115, 928
- Kilgard, R.E., Cowan, J.J., Garcia, M.R., Kaaret, P., Krauss, M.I., McDowell, J.C., et al. 2005, *ApJS*, 159, 214 (K05)
- Kim, D.W., & Fabbiano, G. 2003, *ApJ*, 586, 826 (KF03)
- Kinman, T.D., Green, J.R., & Mahaffey, C.T. 1979, *PASP*, 91, 749
- Knapen, J.H., Beckman, J.E., Cepa, J., & Nakai, N. 1996, *A&A*, 308, 27 (K96)
- Knop, R.A., Armus, L., Matthews, K., Murphy, T.W., & Soifer, B.T. 2001, *ApJ*, 122, 764

- Kohno, K., Ishizuki, S., Matsushita, S., Vila-Vilaró, B., & Kawabe, R. 2003, PASJ, 55, L1 (K03)
- Kohno, K., Tosaki, T., Matsushita, S., Vila-Vilaó, B., Shibatsuka, T., & Kawabe, R. 2002, PASJ, 54, 541 (K02)
- Larkin, J.E., Armus, L., Knop, R.A., Soifer, B.T., & Matthews, K. 1998, ApJS, 114, 59
- Lee, H., Skillman, E.D., Cannon, J.M., et al. 2006, ApJ, 647, 970
- Leitherer, C., Schaerer, D., Goldader, J.D., et al. 1999, ApJS, 123, 3
- Leon, S., Combes, F., & Menon, T.K. 1998, A&A, 330, 37 (L98)
- Lepp, S., & Dalgarno, A. 1996, A&A, 306, L21
- Lepp, S., & McCray, R. 1983, ApJ, 269, 560
- Leroy, A., Bolatto, A.D., Simon, J.D., & Blitz, L. 2005, ApJ, 625, 763 (L05)
- Li, A., & Draine, B.T. 2001, ApJ, 554, 778
- Low, F.J., Young, E., Beintema, D.A., et al. 1984, ApJ, 278, L19
- Lutz, D., Sturm, E., Genzel, R., et al. 2000, ApJ, 536, 697
- Malhotra, S., Kaufman, M.J., Hollenbach, D., et al. 2001, ApJ, 561, 766
- Maloney, P.R., Hollenbach, D.J., & Tielens, A.G. 1996, ApJ, 466, 561
- Maloney, P., & Black, J.H. 1988, ApJ, 325, 389
- Mandy, M.E., & Martin, P.G. 1993, ApJS, 86, 199
- Maoz, D., Koratkar, A., Shields, J.C., Ho, L.C., Filippenko, A.V., & Sternberg, A. 1998, AJ, 116, 55
- Mauersberger, R., Henkel, C., Walsh, W., & Schulz, A. 1999, A&A, 341, 256 (M99)
- Meijerink, R., & Spaans, M. 2005, A&A, 436, 397
- Moneti, A., Stolovy, S., Blommaert, J.A., Figer, D.F., & Najarro, F. 2001, A&A, 366, 106
- Moorwood, A.F., van der Werf, P.P., Kotilainen, J.K., Marconi, A., & Oliva, E. 1996, A&A, 308, L1



- Moorwood, A.F., & Oliva, E. 1988, *A&A*, 203, 278
- Mouri, H., Kawara, K., Taniguchi, Y., & Nishida, M. 1990, *ApJ*, 356, L39
- Nakanishi, H., Sofue, Y., & Koda, J. 2005, *PASJ*, 57, 905
- Nemmen, R.S., Storchi-Bergmann, T., Yuan, F., Eracleous, M., Terashima, Y., & Wilson, A.S. 2006, *ApJ*, 643, 652 (N06)
- Neufeld, D.A., & Kaufman, M.J. 1993, *ApJ*, 418, 263
- Nishiyama, K., & Nakai, N. 2001, *PASJ*, 53, 713 (N01)
- Pagel, B.E., Edmunds, M.G., & Smith, G. 1980, *MNRAS*, 193, 219
- Parmar, P.S., Lacy, J.H., & Achtermann, J.M. 1991, *ApJ*, 372, L25
- Planesas, P., Colina, L., & Perez-Olea, D. 1997, *A&A*, 325, 81 (P97)
- Polk, K.S., Knapp, G.R., Stark, A.A., & Wilson, R.W. 1988, *ApJ*, 332, 432
- Puxley, P.J., Hawarden, T.G., & Mountain, C.M. 1988, *MNRAS*, 234, 29
- Quillen, A.C., Alonso-Herrero, A., Rieke, M.J., Rieke, G.H., Ruiz, M., & Kulkarni, V. 1999, *ApJ*, 527, 696
- Reach, W.T., Rho, J., Jarrett, T.H., & Lagage, P.O. 2002, *ApJ*, 564, 302
- Rieke, G., Young, E.T., Engelbracht, C.W., et al. 2004, *ApJS*, 154, 25
- Rigopoulou, D., Kunze, D., Lutz, D., Genzel, R., & Moorwood, A.F. 2002, *A&A*, 389, 374
- Rotaciuc, V., Krabbe, A., Cameron, M., et al. 1991, *ApJ*, 370, L23
- Roussel, H., Helou, G., Smith, J.D., et al. 2006, *ApJ*, 646, 841
- Sage, L.J. 1993, *A&A*, 272, 123 (S93)
- Sage, L.J., Salzer, J.J., Loose, H.H., & Henkel, C. 1992, *A&A*, 265, 19 (S92)
- Sage, L.J. 1989, *ApJ*, 344, 200 (S89)
- Sakamoto, K., Okumura, S.K., Ishizuki, S., & Scoville, N.Z. 1999, *ApJS*, 124, 403 (S99)
- Satyapal, S., Dudik, R.P., O'Halloran, B., & Gliozzi, M. 2005, *ApJ*, 633, 86 (SD05)
- Satyapal, S., Sambruna, R.M., & Dudik, R.P. 2004, *A&A*, 414, 825 (S04)

- Sellgren, K., Tokunaga, A.T., & Nakada, Y. 1990, *ApJ*, 349, 120
- Sheth, K., Vogel, S.N., Regan, M.W., Thornley, M.D., & Teuben, P.J. 2005, *ApJ*, 632, 217 (S05)
- Shioya, Y., Tosaki, T., Ohyama, Y., et al. 1998, *PASJ*, 50, 317 (S98)
- Shull, J.M., & Beckwith, S. 1982, *ARA&A*, 20, 163
- Shull, J.M., & Hollenbach, D.J. 1978, *ApJ*, 220, 525
- Smith, J.D., Armus, L., Dale, D.A., Roussel, H., Sheth, K., Buckalew, B.A., Helou, G., & Kennicutt, R.C. 2007a, submitted to *PASP*
- Smith, J.D., Draine, B.T., Dale, D.A., et al. 2007b, *ApJ*, 656, 770
- Smith, J.D., Dale, D.A., Armus, L., et al. 2004, *ApJS*, 154, 199
- Sofue, Y., Koda, J., Nakanishi, H., et al. 2003, *PASJ*, 55, 17 (S03)
- Sternberg, A., & Neufeld, D.A. 1999, *ApJ*, 516, 371
- Sternberg, A., & Dalgarno, A. 1989, *ApJ*, 338, 197
- Strong, A.W., Bloemen, J.B., Dame, T.M., et al. 1988, *A&A*, 207, 1
- Sugai, H., Malkan, M.A., Ward, M.J., Davies, R.I., & McLean, I.S. 1997, *ApJ*, 481, 186
- Tacconi, L.J., Tacconi-Garman, L.E., Thornley, M., & van Woerden, H. 1991, *A&A*, 252, 541 (T91)
- Tanaka, M., Hasegawa, T., Hayashi, S.S., Brand, P.W., & Gatley, I. 1989, *ApJ*, 336, 207
- Taniguchi, Y., Murayama, T., Nakai, N., Suzuki, M., & Kameya, O. 1994, *AJ*, 108, 468 (T94)
- Terebey, S., & Fich, M. 1986, *ApJ*, 309, L73
- Thompson, R.I., Lebofsky, M.J., & Rieke, G.H. 1978, *ApJ*, 222, L49
- Tielens, A.G., Meixner, M.M., van der Werf, P.P., et al. 1993, *Science*, 262, 86
- Tielens, A.G., & Hollenbach, D. 1985, *ApJ*, 291, 722
- Timmermann, R. 1998, *ApJ*, 498, 246

- Treffers, R.R. 1979, *ApJ*, 233, L17
- Treffers, R.R., Fink, U., Larson, H.P., & Gautier, T.N. 1976, *ApJ*, 209, 793
- Tschöke, D., Bomans, D.J., Hensler, G. & Junkes, N. 2001, *A&A*, 380, 40
- Usero, A., García-Burillo, S., Fuente, A., Martín-Pintado, J., & Rodríguez-Fernández, N.J. 2004, *A&A*, 419, 897
- Valentijn E.A., van der Werf, P.P., de Graauw, T., & de Jong, T. 1996, *A&A*, 315, L145
- Vila-Vilaró, B., Taniguchi, Y., & Nakai, N. 1998, *AJ*, 116, 1553 (VV98)
- Voit, G.M. 1992, *MNRAS*, 258, 841
- Welch, G.A., & Sage, L.J. 2003, *ApJ*, 584, 260 (W03)
- Wiklind, T., Combes, F., & Henkel, C. 1995, *A&A*, 297, 643 (W95)
- Wilgenbus, D., Cabrit, S., Pineau des Forêts, G., & Flower, D.R. 2000, *A&A*, 356, 1010
- Wu, H., Cao, C., Hao, C.N., et al. 2005, *ApJ*, 632, L79
- Yao, L., Seaquist, E.R., Kuno, N., & Dunne, L. 2003, *ApJ*, 588, 771 (Y03)
- Young, J.S., Xie, S., Tacconi, L., et al. 1995, *ApJS*, 98, 219 (Y95)
- Zhang, X., Wright, M., & Alexander, P. 1993, *ApJ*, 418, 100

Table 1. Observed H<sub>2</sub> lines.<sup>a</sup>

transition v=0	short notation	rest $\lambda$ ( $\mu\text{m}$ )	spectral order	$E_u/k$ (K)	A ( $10^{-11} \text{ s}^{-1}$ )
J=2-0	S(0)	28.219	LH 14	510	2.95
J=3-1	S(1)	17.035	SH 12	1015	47.6
J=4-2	S(2)	12.279	SH 17	1681	275.
J=5-3	S(3)	9.665	SL 1	2503	980.
J=6-4	S(4)	8.025	SL 1	3473	2640.
J=7-5	S(5)	6.910	SL 2	4585	5880.
J=8-6	S(6)	6.109	SL 2	5828	11400.
J=9-7	S(7)	5.511	SL 2	7196	20000.

<sup>a</sup>The rotational upper level energies were computed from the molecular constants given by Huber & Herzberg (1979) and the transition probabilities are from Black & Dalgarno (1976).

Table 2. Targets.<sup>a</sup>

galaxy	D (Mpc)	class	center (J2000)		solid angle (arcsec <sup>2</sup> )
			RA	DEC	
N24	8.2	dwarf	00 09 56.31	−24 57 52.0	294.
N337	24.7	nuc	00 59 49.99	−07 34 42.5	289.
N628	11.4	nuc	01 36 41.72	+15 46 59.4	300.
N855	9.6	dwarf	02 14 03.64	+27 52 40.6	288.
N925	10.1	nuc	02 27 17.06	+33 34 43.8	287.
N1097	16.9	nuc	02 46 18.77	−30 16 30.0	809.
N1266	31.3	liner	03 16 00.68	−02 25 38.8	287.
N1291	9.7	liner	03 17 18.50	−41 06 27.7	287.
N1316	26.3	liner	03 22 41.61	−37 12 28.7	287.
N1482	22.0	nuc	03 54 38.68	−20 30 08.5	798.
N1512	10.4	nuc	04 03 53.96	−43 20 55.6	831.
N1566	18.0	sy	04 20 00.53	−54 56 17.4	306.
N1705	5.8	dwarf	04 54 13.26	−53 21 39.2	802.
N2403	3.5	nuc	07 36 50.25	+65 36 04.6	287.
N2798	24.7	nuc	09 17 22.99	+42 00 00.8	293.
N2841	9.8	liner	09 22 02.75	+50 58 35.5	291.
N2915	2.7	dwarf	09 26 11.59	−76 37 34.2	291.
N2976	3.5	dwarf	09 47 15.63	+67 55 00.3	289.
N3031	3.5	sy	09 55 33.50	+69 03 55.8	296.
N3049	19.6	nuc	09 54 49.66	+09 16 19.2	320.
N3184	8.6	nuc	10 18 17.03	+41 25 28.3	287.
N3190	17.4	liner	10 18 05.71	+21 49 56.8	304.
N3198	9.8	nuc	10 19 55.10	+45 32 59.6	287.
N3265	20.0	nuc	10 31 06.85	+28 47 49.1	301.
Mrk33	21.7	dwarf	10 32 32.02	+54 24 05.4	315.
N3351	9.3	nuc	10 43 57.83	+11 42 10.9	761.
N3521	9.0	liner	11 05 48.51	−00 02 10.1	285.
N3621	6.2	liner	11 18 16.62	−32 48 49.9	291.
N3627	8.9	sy	11 20 15.04	+12 59 31.0	301.
N3773	12.9	dwarf	11 38 12.91	+12 06 43.0	285.
N3938	12.2	nuc	11 52 49.45	+44 07 15.0	292.
N4125	21.4	liner	12 08 05.73	+65 10 27.9	294.
N4254	20.0	nuc	12 18 49.71	+14 25 01.4	299.
N4321	20.0	nuc	12 22 55.08	+15 49 18.2	795.
N4450	20.0	liner	12 28 29.51	+17 05 04.9	285.

Table 2—Continued

galaxy	D (Mpc)	class	center (J2000)		solid angle (arcsec <sup>2</sup> )
			RA	DEC	
N4536	25.0	nuc	12 34 27.22	+02 11 14.3	777.
N4552	20.0	liner	12 35 39.71	+12 33 21.9	285.
N4559	11.6	nuc	12 35 57.79	+27 57 36.7	295.
N4569	20.0	liner	12 36 49.91	+13 09 46.8	287.
N4579	20.0	sy	12 37 43.70	+11 49 07.1	295.
N4594	13.7	liner	12 39 59.58	−11 37 22.3	289.
N4625	9.5	dwarf	12 41 52.65	+41 16 26.5	294.
N4631	9.0	nuc	12 42 07.84	+32 32 34.8	282.
N4725	17.1	sy	12 50 26.71	+25 30 03.2	294.
N4736	5.3	liner	12 50 53.23	+41 07 13.6	298.
N4826	5.6	liner	12 56 43.79	+21 41 00.7	292.
N5033	13.3	sy	13 13 27.67	+36 35 38.2	311.
N5055	8.2	liner	13 15 49.50	+42 01 46.6	313.
N5194	8.2	sy	13 29 52.94	+47 11 44.2	330.
N5194_A <sup>b</sup>	8.2	hii	13 29 49.54	+47 13 28.6	256.
N5194_B <sup>b</sup>	8.2	hii	13 30 01.69	+47 12 50.9	271.
N5194_C <sup>b</sup>	8.2	hii	13 30 00.02	+47 11 11.9	352.
N5194_D <sup>b</sup>	8.2	hii	13 30 02.76	+47 09 53.0	338.
N5194_E <sup>b</sup>	8.2	hii	13 29 56.84	+47 10 45.9	234.
N5194_F <sup>b</sup>	8.2	hii	13 29 52.93	+47 12 38.7	297.
N5194_G <sup>b</sup>	8.2	hii	13 29 43.99	+47 10 20.7	285.
N5195	8.2	liner	13 29 59.80	+47 16 00.1	323.
N5713	26.6	nuc	14 40 11.57	−00 17 19.1	292.
N5866	12.5	liner	15 06 29.58	+55 45 46.7	286.
N6822_A <sup>c</sup>	0.6	dwarf	19 44 52.84	−14 43 09.8	308.
N6822_B <sup>c</sup>	0.6	dwarf	19 44 50.57	−14 52 49.0	288.
N6822_C <sup>c</sup>	0.6	dwarf	19 44 48.67	−14 52 26.5	304.
N6946	5.5	nuc	20 34 51.79	+60 09 12.6	803.
N7331	15.7	liner	22 37 04.11	+34 24 57.5	263.
N7552	22.3	nuc	23 16 10.56	−42 35 04.0	722.
N7793	3.2	dwarf	23 57 49.79	−32 35 28.4	296.

<sup>a</sup>Distances are from Kennicutt et al. (2003). Nuclear classifications derived from optical spectroscopy, published by Smith et al. (2007b), were modified for the following targets. N1097, N1512, N4321 and N7552: Since our aperture includes a bright star-forming ring, which dominates the dust and line emission,

we adopt the HII class instead of the LINER class. N2841, N4552, N4569, N4594, N4826 and N5195: The Sy class was changed to LINER (Ho et al. 1997). N3198 and N3938: The LINER class was changed to HII (Ho et al. 1997). The dwarf galaxy class is here arbitrarily defined by a total stellar mass, estimated as by Lee et al. (2006), below  $10^{9.7} M_{\odot}$ .

<sup>b</sup>A total of 11 extranuclear regions were observed in high-resolution spectroscopy in NGC 5194, from which we included 7 in our sample.

<sup>c</sup>Several locations were observed in high-resolution spectroscopy in this galaxy. We present here results for three among the brightest regions at 7.9 and  $24 \mu\text{m}$ . N6822\_A corresponds to the HII region Hubble V (Hubble 1925), and N6822\_C to the HII region K $\gamma$  (Kinman et al. 1979).

Table 3. Fluxes.<sup>a</sup>

galaxy	S(0)	S(1)	S(2) ( $10^{-18} \text{ W m}^{-2}$ )	S(3)	[SiII]	$F_{7.9 \text{ dust}}$ (Jy)	$F_{24}$ (Jy)	TIR ( $10^{-15} \text{ W m}^{-2}$ )
N24	$6.6 \pm 2.6$	$2.2 \pm 2.4$	$3.2 \pm 2.0$	$10.3 \pm 6.4$	$36.7 \pm 3.6$	0.010	0.010	$18. \pm 3.$
N337	$11.6 \pm 3.8$	$21.0 \pm 4.9$	$15.4 \pm 9.9$	$9.7 \pm 5.2$	$186.1 \pm 7.0$	0.048	0.087	$89. \pm 9.$
N628	$7.9 \pm 2.5$	$13.7 \pm 5.3$	$0.0 \pm 3.0$	$5.7 \pm 8.6$	$47.8 \pm 2.8$	0.026	0.026	$39. \pm 3.$
N855	$2.9 \pm 0.9$	$11.2 \pm 6.5$	$0.0 \pm 5.1$	$22.4 \pm 7.2$	$70.2 \pm 3.8$	0.017	0.039	$53. \pm 7.$
N925	$3.8 \pm 1.7$	$7.8 \pm 2.9$	$3.7 \pm 2.1$	$0.0 \pm 8.8$	$81.0 \pm 4.5$	0.021	0.022	$37. \pm 4.$
N1097	$213.1 \pm 42.6$	$726.1 \pm 43.3$	$293.6 \pm 29.4$	$423.0 \pm 23.4$	$4921.1 \pm 121.9$	1.139	3.763	$2214. \pm 227.$
N1266 <sup>b</sup>	$16.4 \pm 8.5$	$148.5 \pm 6.6$	$121.8 \pm 7.1$	$189.8 \pm 11.6$	$153.5 \pm 41.5$	0.048	0.579	$438. \pm 30.$
N1291	$3.7 \pm 1.8$	$29.8 \pm 4.8$	$13.9 \pm 8.2$	$31.6 \pm 10.0$	$75.1 \pm 5.8$	0.013	0.049	$60. \pm 7.$
N1316	$1.5 \pm 0.8$	$36.0 \pm 6.1$	$18.4 \pm 5.9$	$84.0 \pm 9.9$	$102.4 \pm 6.4$	0.025	0.094	$105. \pm 11.$
N1482	$106.8 \pm 35.1$	$424.0 \pm 53.4$	$184.0 \pm 22.8$	$207.5 \pm 26.9$	$3954.0 \pm 75.1$	1.163	2.974	$1752. \pm 154.$
N1512	$25.9 \pm 2.9$	$83.6 \pm 8.9$	$27.2 \pm 10.4$	$48.3 \pm 13.6$	$390.7 \pm 13.9$	0.091	0.150	$203. \pm 24.$
N1566	$24.0 \pm 2.2$	$129.5 \pm 7.1$	$55.3 \pm 4.8$	$91.6 \pm 17.6$	$101.9 \pm 7.8$	0.098	0.237	$235. \pm 19.$
N1705	$5.6 \pm 3.9$	$2.2 \pm 1.3$	$0.0 \pm 10.3$	$11.6 \pm 4.0$	$86.3 \pm 12.0$	0.007	0.028	$40. \pm 5.$
N2403	$5.5 \pm 1.5$	$6.6 \pm 3.1$	$2.0 \pm 2.0$	$4.6 \pm 8.2$	$82.0 \pm 5.9$	0.029	0.022	$35. \pm 2.$
N2798	$43.2 \pm 25.3$	$207.8 \pm 22.1$	$90.7 \pm 9.2$	$105.2 \pm 12.1$	$840.4 \pm 42.2$	0.354	1.723	$936. \pm 67.$
N2841	$3.1 \pm 1.4$	$9.6 \pm 2.5$	$4.1 \pm 4.5$	$28.0 \pm 10.8$	$81.1 \pm 5.5$	0.010	0.027	$48. \pm 4.$
N2915	$0.8 \pm 0.6$	$1.8 \pm 2.1$	$3.8 \pm 1.8$	$0.0 \pm 8.0$	$40.1 \pm 3.7$	0.005	0.019	$25. \pm 3.$
N2976	$8.9 \pm 2.4$	$15.7 \pm 3.0$	$4.9 \pm 3.5$	$5.2 \pm 4.2$	$66.1 \pm 5.4$	0.026	0.055	$77. \pm 6.$
N3031	$6.3 \pm 3.6$	$46.1 \pm 6.1$	$26.1 \pm 8.1$	$65.3 \pm 8.0$	$217.9 \pm 9.7$	0.081	0.337	$234. \pm 17.$
N3049	$6.4 \pm 2.3$	$23.7 \pm 3.8$	$11.7 \pm 6.5$	$11.1 \pm 5.2$	$197.1 \pm 11.9$	0.051	0.258	$141. \pm 13.$
N3184	$9.8 \pm 2.3$	$22.6 \pm 2.9$	$6.3 \pm 3.5$	$0.0 \pm 15.9$	$105.0 \pm 4.2$	0.031	0.082	$103. \pm 13.$
N3190	$18.1 \pm 2.9$	$75.3 \pm 8.0$	$20.9 \pm 6.4$	$71.6 \pm 12.6$	$67.1 \pm 7.3$	0.051	0.068	$137. \pm 23.$
N3198	$13.2 \pm 4.9$	$32.1 \pm 5.4$	$10.2 \pm 3.5$	$18.7 \pm 8.5$	$96.5 \pm 6.6$	0.059	0.260	$200. \pm 20.$
N3265	$9.3 \pm 3.2$	$30.6 \pm 3.5$	$11.4 \pm 7.9$	$15.3 \pm 7.4$	$159.1 \pm 7.7$	0.062	0.193	$119. \pm 9.$
Mrk33	$13.9 \pm 5.6$	$29.3 \pm 3.7$	$11.5 \pm 5.3$	$45.7 \pm 8.6$	$245.4 \pm 10.9$	0.076	0.583	$239. \pm 13.$
N3351	$62.8 \pm 10.5$	$218.4 \pm 20.6$	$88.3 \pm 10.1$	$169.8 \pm 19.9$	$1858.9 \pm 35.7$	0.366	1.385	$917. \pm 93.$



Table 3—Continued

galaxy	S(0)	S(1)	S(2) ( $10^{-18} \text{ W m}^{-2}$ )	S(3)	[SiIII]	$F_{7.9 \text{ dust}}$ (Jy)	$F_{24}$ (Jy)	TIR ( $10^{-15} \text{ W m}^{-2}$ )
N3521	$17.8 \pm 3.6$	$31.2 \pm 6.9$	$12.3 \pm 4.0$	$20.5 \pm 8.8$	$157.5 \pm 4.6$	0.101	0.106	$129. \pm 10.$
N3621	$16.7 \pm 2.1$	$38.3 \pm 3.2$	$17.1 \pm 4.7$	$49.0 \pm 8.4$	$150.3 \pm 7.5$	0.075	0.065	$88. \pm 6.$
N3627	$31.2 \pm 3.9$	$318.6 \pm 12.2$	$141.7 \pm 5.2$	$209.1 \pm 12.6$	$162.1 \pm 9.3$	0.190	0.406	$518. \pm 31.$
N3773	$4.8 \pm 3.4$	$7.3 \pm 3.2$	$3.7 \pm 2.5$	$8.3 \pm 8.6$	$106.9 \pm 6.1$	0.024	0.087	$67. \pm 7.$
N3938	$8.0 \pm 1.2$	$13.4 \pm 3.3$	$4.2 \pm 3.2$	$0.0 \pm 8.0$	$57.7 \pm 4.2$	0.027	0.030	$51. \pm 5.$
N4125	$2.6 \pm 1.6$	$17.6 \pm 4.7$	$12.6 \pm 6.4$	$18.0 \pm 8.8$	$46.8 \pm 5.9$	0.005	0.015	$21. \pm 3.$
N4254	$17.6 \pm 1.3$	$87.3 \pm 9.5$	$46.6 \pm 9.2$	$32.5 \pm 8.8$	$409.0 \pm 10.6$	0.160	0.203	$231. \pm 23.$
N4321	$78.2 \pm 11.7$	$267.5 \pm 22.8$	$132.1 \pm 24.6$	$164.5 \pm 16.2$	$1507.3 \pm 27.2$	0.392	0.730	$778. \pm 90.$
N4450	$10.3 \pm 1.6$	$91.4 \pm 5.2$	$33.7 \pm 9.5$	$89.0 \pm 11.2$	$29.6 \pm 3.1$	0.011	0.028	$55. \pm 9.$
N4536	$108.7 \pm 32.3$	$411.6 \pm 40.5$	$175.6 \pm 23.3$	$216.8 \pm 20.7$	$3207.4 \pm 70.0$	0.813	2.401	$1400. \pm 118.$
N4552	$1.8 \pm 1.4$	$1.1 \pm 1.5$	$0.0 \pm 3.9$	$9.1 \pm 4.6$	$13.4 \pm 3.9$	0.004	0.017	$6. \pm 1.$
N4559	$13.6 \pm 2.1$	$17.9 \pm 3.5$	$4.0 \pm 1.5$	$23.4 \pm 13.7$	$81.1 \pm 3.5$	0.034	0.030	$47. \pm 5.$
N4569 <sup>c</sup>	$41.6 \pm 6.9$	$316.1 \pm 8.7$	$152.4 \pm 5.2$	$309.4 \pm 18.7$	$288.5 \pm 17.6$	0.128	0.475	$365. \pm 40.$
N4579 <sup>d</sup>	$8.9 \pm 2.3$	$164.1 \pm 5.2$	$105.8 \pm 12.4$	$256.3 \pm 10.0$	$175.3 \pm 12.2$	0.032	0.135	$160. \pm 18.$
N4594	$1.9 \pm 2.3$	$6.4 \pm 3.8$	$4.2 \pm 2.7$	$4.6 \pm 4.0$	$126.1 \pm 7.3$	0.017	0.068	$70. \pm 8.$
N4625	$8.0 \pm 1.6$	$9.6 \pm 1.9$	$5.4 \pm 4.3$	$5.0 \pm 11.4$	$67.5 \pm 4.8$	0.024	0.024	$33. \pm 5.$
N4631	$58.2 \pm 6.4$	$123.6 \pm 16.1$	$60.4 \pm 5.6$	$46.8 \pm 9.1$	$1350.4 \pm 22.7$	0.338	0.661	$636. \pm 61.$
N4725	$11.7 \pm 1.7$	$37.1 \pm 5.1$	$19.9 \pm 8.1$	$38.0 \pm 9.0$	$24.9 \pm 3.2$	0.011	0.035	$56. \pm 9.$
N4736	$33.2 \pm 6.8$	$250.8 \pm 16.4$	$100.1 \pm 10.4$	$221.7 \pm 13.6$	$219.5 \pm 21.4$	0.293	0.636	$678. \pm 50.$
N4826	$71.3 \pm 6.5$	$345.3 \pm 15.8$	$150.9 \pm 10.9$	$210.6 \pm 10.3$	$927.6 \pm 15.0$	0.343	0.612	$756. \pm 92.$
N5033	$36.6 \pm 3.5$	$182.0 \pm 10.4$	$63.5 \pm 3.1$	$126.9 \pm 19.1$	$530.9 \pm 10.0$	0.177	0.239	$311. \pm 45.$
N5055	$44.0 \pm 2.4$	$158.0 \pm 9.1$	$52.0 \pm 6.7$	$80.2 \pm 10.6$	$295.2 \pm 9.0$	0.148	0.166	$300. \pm 37.$
N5194	$17.7 \pm 2.8$	$134.4 \pm 9.1$	$75.7 \pm 4.7$	$187.3 \pm 17.7$	$634.9 \pm 12.2$	0.168	0.291	$351. \pm 39.$
N5194_A	$5.8 \pm 1.3$	$16.4 \pm 2.7$	$4.3 \pm 2.6$	$5.4 \pm 1.8$	$27.8 \pm 3.0$	0.030	0.036	$53. \pm 2.$
N5194_B	$18.1 \pm 2.2$	$43.1 \pm 8.2$	$12.0 \pm 4.3$	$18.1 \pm 10.2$	$339.7 \pm 8.5$	0.094	0.193	$201. \pm 13.$
N5194_C	$23.2 \pm 2.2$	$43.4 \pm 6.7$	$5.7 \pm 4.5$	$14.1 \pm 4.6$	$223.6 \pm 6.8$	0.079	0.115	$155. \pm 7.$

Table 3—Continued

galaxy	S(0)	S(1)	S(2) ( $10^{-18} \text{ W m}^{-2}$ )	S(3)	[SiII]	$F_{7.9 \text{ dust}}$ (Jy)	$F_{24}$ (Jy)	TIR ( $10^{-15} \text{ W m}^{-2}$ )
N5194_D	$14.7 \pm 2.8$	$30.3 \pm 6.5$	$3.9 \pm 3.1$	$8.4 \pm 3.0$	$222.9 \pm 6.2$	0.058	0.118	$130. \pm 14.$
N5194_E	$22.7 \pm 1.7$	$47.9 \pm 6.0$	$19.2 \pm 7.1$	$23.3 \pm 3.1$	$222.2 \pm 4.8$	0.072	0.089	$116. \pm 8.$
N5194_F	$19.4 \pm 2.2$	$39.2 \pm 5.2$	$15.4 \pm 9.1$	$21.6 \pm 14.0$	$271.7 \pm 5.8$	0.075	0.110	$132. \pm 4.$
N5194_G	$22.2 \pm 3.0$	$49.6 \pm 5.4$	$23.6 \pm 2.7$	$27.7 \pm 11.2$	$425.6 \pm 8.5$	0.093	0.192	$173. \pm 15.$
N5195	$54.2 \pm 20.9$	$310.4 \pm 12.3$	$127.7 \pm 6.5$	$275.0 \pm 17.2$	$127.8 \pm 19.6$	0.271	0.835	$600. \pm 55.$
N5713	$27.4 \pm 3.6$	$151.6 \pm 14.3$	$50.9 \pm 7.9$	$85.6 \pm 7.0$	$798.7 \pm 14.8$	0.268	0.738	$500. \pm 54.$
N5866	$15.5 \pm 1.5$	$90.0 \pm 5.5$	$39.1 \pm 7.0$	$37.2 \pm 6.5$	$78.8 \pm 3.9$	0.038	0.052	$169. \pm 26.$
N6822_A	$8.1 \pm 4.3$	$21.1 \pm 4.3$	$16.0 \pm 4.1$	$25.9 \pm 3.8$	$187.1 \pm 11.2$	0.036	0.517	$283. \pm 24.$
N6822_B	$7.8 \pm 1.3$	$11.2 \pm 3.6$	$8.8 \pm 2.8$	$10.9 \pm 5.1$	$19.0 \pm 2.8$	0.010	0.037	$44. \pm 5.$
N6822_C	$4.4 \pm 1.9$	$8.0 \pm 5.2$	$6.2 \pm 5.4$	$8.0 \pm 2.9$	$16.0 \pm 3.6$	0.009	0.100	$76. \pm 7.$
N6946	$227.7 \pm 61.9$	$636.9 \pm 33.5$	$271.6 \pm 17.1$	$286.3 \pm 29.3$	$3203.9 \pm 91.9$	1.314	4.234	$2354. \pm 193.$
N7331	$16.4 \pm 1.6$	$49.6 \pm 5.0$	$13.8 \pm 1.9$	$23.4 \pm 8.3$	$213.9 \pm 4.8$	0.074	0.091	$125. \pm 10.$
N7552	$257.8 \pm 106.9$	$565.3 \pm 42.6$	$245.9 \pm 22.5$	$315.9 \pm 14.8$	$3619.9 \pm 165.2$	1.495	7.840	$3237. \pm 219.$
N7793	$7.0 \pm 1.7$	$15.0 \pm 3.2$	$11.7 \pm 4.6$	$5.3 \pm 3.7$	$79.7 \pm 4.3$	0.041	0.046	$71. \pm 5.$

<sup>a</sup>See Sections 2.2 and 2.4. For NGC 1266, NGC 4569 and NGC 4579, it was possible to measure higher-level transitions of H<sub>2</sub>, which are provided in the following notes.

<sup>b</sup>The S(4) to S(7) line fluxes are respectively: ( $102.7 \pm 17.4$ ), ( $241.9 \pm 19.9$ ),  $> 18.5$  and ( $192.2 \pm 28.6$ )  $\times 10^{-18} \text{ W m}^{-2}$ . The S(6) transition was not accurately measurable because it was observed at low spectral resolution on the blue shoulder of a bright  $6.2 \mu\text{m}$  aromatic band.

<sup>c</sup>The S(4) to S(7) line fluxes are respectively: ( $52.6 \pm 27.7$ ),  $< 252.6$ ,  $< 61.3$  and ( $50.3 \pm 24.2$ )  $\times 10^{-18} \text{ W m}^{-2}$ . The S(5) and S(6) transitions are upper limits because S(5) is contaminated by the [ArII] line at  $6.985 \mu\text{m}$  and S(6) is diluted in the blue shoulder of the  $6.2 \mu\text{m}$  aromatic band.

<sup>d</sup>The S(4) to S(7) line fluxes are respectively:  $(48.7 \pm 10.7)$ ,  $< 324.9$ ,  $< 147.2$  and  $(83.3 \pm 37.5) \times 10^{-18} \text{ W m}^{-2}$ . The same remark as for NGC 4569 applies.

Table 4. Derived temperatures, column densities and  $OPR_{\text{high T}}$ .<sup>a</sup>

galaxy	$T_1$ (K)	$f_1$	$T_2$ (K)	$f_2$	$OPR_{\text{high T}}$	$N_{\text{tot}}(T > T_1)$ ( $10^{20}$ mol. cm $^{-2}$ )	$N_{\text{tot}}(OPR = 3)$ $/N_{\text{tot}}(OPR < 3)$
N24	97.	0.99941	769.	0.00059	$1.49 \pm 0.91$	4.232	1.75
	78.	0.99980	675.	0.00020	$1.90 \pm 0.57$	15.210	1.46
	90.	0.99971	888.	0.00029	$1.45 \pm 0.37$	7.314	1.73
N337	160.	0.97793	402.	0.02207	$1.02 \pm 0.15$	1.745	2.13
	128.	0.97803	352.	0.02197	$1.21 \pm 0.25$	3.218	0.97
	165.	0.98711	464.	0.01289	$0.86 \pm 0.14$	1.824	2.19
N628	119.	0.99678	393.	0.00322	3	2.598	
N855	146.	0.98750	578.	0.01250	3	0.508	
N925	123.	0.99473	494.	0.00527	3	1.050	
	99.	0.99648	421.	0.00352	3	2.780	
N1097	157.	0.98027	422.	0.01973	$2.03 \pm 0.36$	11.837	1.29
	94.	0.99033	335.	0.00967	3	62.925	
N1266 <sup>b</sup>	247.	0.83691	400.	0.15619	$1.79 \pm 0.29$	1.231	1.19
	204.	0.82588	400.	0.16935	$1.80 \pm 0.18$	1.594	0.84
N1291	184.	0.95420	484.	0.04580	3	0.415	
	141.	0.95010	414.	0.04990	3	0.762	
N1316	300.	0.96436	1456.	0.03564	3	0.114	
	> 278.	> 0.93750	> 900.	< 0.06250	3	0.121	
N1482	171.	0.96943	407.	0.03057	$1.83 \pm 0.25$	5.098	1.38
	137.	0.96699	357.	0.03301	$2.12 \pm 0.43$	8.848	1.49
	175.	0.98252	469.	0.01748	$1.58 \pm 0.25$	5.378	1.40
N1512	138.	0.98887	415.	0.01113	3	1.883	
	118.	0.98955	375.	0.01045	3	3.338	
N1566	177.	0.96436	441.	0.03564	$2.24 \pm 0.37$	2.749	1.19
	115.	0.97305	359.	0.02695	3	8.517	
N1705	127.	0.99453	811.	0.00547	$0.49 \pm 0.29$	0.494	3.78
	92.	0.99844	711.	0.00156	$0.65 \pm 0.14$	2.050	0.95
	114.	0.99727	936.	0.00273	$0.49 \pm 0.11$	0.876	2.33
N2403	109.	0.99844	449.	0.00156	3	2.486	
	99.	0.99873	419.	0.00127	3	4.138	
N2798	179.	0.96182	409.	0.03818	$1.92 \pm 0.23$	5.169	1.32
	144.	0.95518	359.	0.04482	$2.17 \pm 0.43$	8.460	1.16
	184.	0.97812	472.	0.02188	$1.67 \pm 0.27$	5.419	1.27
N2841	137.	0.99678	1030.	0.00322	3	0.709	
	132.	0.99707	970.	0.00293	3	0.844	

Table 4—Continued

galaxy	$T_1$ (K)	$f_1$	$T_2$ (K)	$f_2$	$OPR_{\text{high T}}$	$N_{\text{tot}}(T > T_1)$ ( $10^{20}$ mol. cm $^{-2}$ )	$N_{\text{tot}}(OPR = 3)$ $/N_{\text{tot}}(OPR < 3)$
N2915	169.	0.97090	763.	0.02910	$1.05 \pm 0.47$	0.096	1.74
	116.	0.98838	669.	0.01162	$1.30 \pm 0.28$	0.305	0.55
	151.	0.98350	880.	0.01650	$1.00 \pm 0.20$	0.150	1.13
N2976	134.	0.99199	376.	0.00801	$1.84 \pm 0.30$	2.095	1.43
	86.	0.99619	300.	0.00381	3	12.046	
N3031	178.	0.95908	523.	0.04092	3	0.700	
	126.	0.96387	433.	0.03613	3	1.699	
N3049	174.	0.96611	403.	0.03389	$1.60 \pm 0.20$	0.740	1.50
	141.	0.96113	353.	0.03887	$1.83 \pm 0.36$	1.238	1.17
	179.	0.98076	465.	0.01924	$1.39 \pm 0.22$	0.773	1.55
N3184	127.	0.99502	440.	0.00498	3	2.631	
	118.	0.99648	449.	0.00352	3	3.805	
N3190	148.	0.99355	574.	0.00645	3	3.286	
	142.	0.99385	554.	0.00615	3	3.883	
N3198	128.	0.99316	415.	0.00684	3	3.438	
	111.	0.99385	375.	0.00615	3	6.283	
N3265	156.	0.98105	408.	0.01895	$2.02 \pm 0.35$	1.421	1.30
	95.	0.98975	325.	0.01025	3	7.006	
Mrk33	124.	0.99736	624.	0.00264	3	3.844	
	115.	0.99775	584.	0.00225	3	5.478	
N3351	141.	0.98633	437.	0.01367	3	4.575	
	114.	0.98867	387.	0.01133	3	9.893	
N3521	118.	0.99531	414.	0.00469	3	5.715	
	91.	0.99687	354.	0.00313	3	18.837	
N3621	126.	0.99541	532.	0.00459	3	4.490	
	110.	0.99609	472.	0.00391	3	8.195	
N3627	220.	0.91836	433.	0.08164	$2.42 \pm 0.19$	2.672	1.12
	124.	0.90469	344.	0.09531	3	7.324	
N3773	115.	0.99678	480.	0.00322	3	1.671	
	90.	0.99805	410.	0.00195	3	5.446	
N3938	118.	0.99629	425.	0.00371	3	2.603	
	102.	0.99697	377.	0.00303	3	5.116	
N4125	211.	0.92803	481.	0.07197	$1.75 \pm 0.24$	0.238	1.35
	163.	0.91826	422.	0.08174	$1.90 \pm 0.37$	0.377	1.15
	212.	0.95625	555.	0.04375	$1.56 \pm 0.25$	0.262	1.35

Table 4—Continued

galaxy	$T_1$ (K)	$f_1$	$T_2$ (K)	$f_2$	$OPR_{\text{high T}}$	$N_{\text{tot}}(T > T_1)$ ( $10^{20}$ mol. cm $^{-2}$ )	$N_{\text{tot}}(OPR = 3)$ $/N_{\text{tot}}(OPR < 3)$
N4254	200.	0.94111	384.	0.05889	$1.49 \pm 0.11$	1.763	1.53
	167.	0.90293	337.	0.09707	$1.59 \pm 0.29$	2.368	1.14
	206.	0.96982	444.	0.03018	$1.36 \pm 0.22$	1.812	1.45
N4321	164.	0.97480	427.	0.02520	$1.76 \pm 0.29$	4.001	1.41
	131.	0.97627	374.	0.02373	$2.07 \pm 0.43$	7.511	1.27
	168.	0.98477	493.	0.01523	$1.49 \pm 0.25$	4.258	1.46
N4450	190.	0.96992	536.	0.03008	3	1.178	
	172.	0.96152	476.	0.03848	3	1.467	
N4536	166.	0.97344	413.	0.02656	$1.90 \pm 0.28$	5.602	1.35
	133.	0.97334	362.	0.02666	$2.23 \pm 0.45$	10.170	1.35
	170.	0.98437	476.	0.01563	$1.62 \pm 0.27$	5.908	1.40
N4552	126.	0.99463	986.	0.00537	$0.84 \pm 0.56$	0.449	2.61
	89.	0.99873	865.	0.00127	$1.09 \pm 0.24$	2.156	0.54
	108.	0.99766	1138.	0.00234	$0.87 \pm 0.19$	0.959	1.30
N4559	111.	0.99951	775.	0.00049	3	6.202	
	109.	0.99961	825.	0.00039	3	6.967	
N4569 <sup>c</sup>	184.	0.91348	400.	0.08564	3	4.062	
	138.	0.95226	400.	0.04755	3	9.055	
N4579 <sup>d</sup>	286.	0.55439	400.	0.43999	$2.80 \pm 0.56$	0.534	1.00
	68.	0.98551	400.	0.01439	3	21.592	
N4594	175.	0.96621	441.	0.03379	$1.49 \pm 0.23$	0.232	1.57
	138.	0.96709	387.	0.03291	$1.71 \pm 0.35$	0.421	1.06
	178.	0.97969	509.	0.02031	$1.28 \pm 0.21$	0.251	1.64
N4625	133.	0.99268	411.	0.00732	$1.27 \pm 0.28$	1.856	1.85
	107.	0.99453	361.	0.00547	$1.65 \pm 0.35$	4.144	1.20
	136.	0.99551	475.	0.00449	$1.05 \pm 0.20$	1.964	2.38
N4631	154.	0.98174	386.	0.01826	$1.36 \pm 0.19$	9.852	1.75
	125.	0.98174	338.	0.01826	$1.64 \pm 0.32$	18.343	0.96
	159.	0.98955	445.	0.01045	$1.14 \pm 0.19$	10.160	2.01
N4725	155.	0.98223	481.	0.01777	$1.97 \pm 0.50$	1.814	1.31
	88.	0.99482	381.	0.00518	3	13.562	
N4736	180.	0.96455	478.	0.03545	3	3.822	
	151.	0.95732	418.	0.04268	3	5.791	
N4826	176.	0.96523	426.	0.03477	$2.05 \pm 0.30$	8.766	1.27
	72.	0.99570	328.	0.00430	3	223.50	

Table 4—Continued

galaxy	$T_1$ (K)	$f_1$	$T_2$ (K)	$f_2$	$OPR_{\text{high T}}$	$N_{\text{tot}}(T > T_1)$ ( $10^{20}$ mol. cm $^{-2}$ )	$N_{\text{tot}}(OPR = 3)$ $/N_{\text{tot}}(OPR < 3)$
N5033	157.	0.97900	438.	0.02100	3	5.255	
	134.	0.97920	398.	0.02080	3	8.435	
N5055	142.	0.98242	386.	0.01758	3	7.586	
	114.	0.98506	346.	0.01494	3	16.336	
N5194	180.	0.95732	521.	0.04268	3	1.753	
	127.	0.96123	431.	0.03877	3	4.171	
N5194_A	133.	0.98672	349.	0.01328	3	1.469	
	111.	0.98848	319.	0.01152	3	2.951	
N5194_B	128.	0.99150	373.	0.00850	3	4.933	
	109.	0.99287	343.	0.00713	3	9.421	
N5194_C	120.	0.99873	458.	0.00127	3	6.824	
	118.	0.99873	448.	0.00127	3	7.448	
N5194_D	123.	0.99824	429.	0.00176	3	4.144	
	121.	0.99824	419.	0.00176	3	4.572	
N5194_E	142.	0.98896	406.	0.01104	$1.79 \pm 0.33$	5.653	1.44
	63.	0.99922	313.	0.00078	3	235.5	
N5194_F	137.	0.99102	417.	0.00898	$1.91 \pm 0.40$	4.127	1.37
	85.	0.99648	331.	0.00352	3	26.208	
N5194_G	148.	0.98594	418.	0.01406	$1.64 \pm 0.31$	4.041	1.53
	118.	0.98857	367.	0.01143	$2.03 \pm 0.43$	8.457	1.21
	151.	0.99160	483.	0.00840	$1.37 \pm 0.24$	4.337	1.74
N5195	164.	0.97441	468.	0.02559	3	6.769	
	134.	0.97305	408.	0.02695	3	11.899	
N5713	162.	0.97266	413.	0.02734	3	3.901	
	133.	0.96875	363.	0.03125	3	6.684	
N5866	194.	0.94678	393.	0.05322	$1.88 \pm 0.15$	1.676	1.33
	160.	0.92070	344.	0.07930	$2.03 \pm 0.37$	2.379	1.10
	200.	0.97158	453.	0.02842	$1.68 \pm 0.27$	1.731	1.25
N6822_A	158.	0.98027	504.	0.01973	$1.53 \pm 0.42$	1.094	1.56
	120.	0.98721	442.	0.01279	$1.89 \pm 0.41$	2.561	1.31
	156.	0.98799	581.	0.01201	$1.33 \pm 0.25$	1.306	1.76
N6822_B	140.	0.99043	476.	0.00957	$1.29 \pm 0.35$	1.609	1.82
	109.	0.99404	417.	0.00596	$1.63 \pm 0.36$	3.946	1.30
	139.	0.99414	549.	0.00586	$1.09 \pm 0.21$	1.867	2.00
N6822_C	148.	0.98604	478.	0.01396	$1.33 \pm 0.35$	0.716	1.77

Table 4—Continued

galaxy	$T_1$ (K)	$f_1$	$T_2$ (K)	$f_2$	$OPR_{\text{high T}}$	$N_{\text{tot}}(T > T_1)$ ( $10^{20}$ mol. cm $^{-2}$ )	$N_{\text{tot}}(OPR = 3)$ $/N_{\text{tot}}(OPR < 3)$
	115.	0.99092	419.	0.00908	$1.67 \pm 0.34$	1.690	1.04
	148.	0.99150	552.	0.00850	$1.14 \pm 0.21$	0.830	2.19
N6946	156.	0.98066	399.	0.01934	$1.70 \pm 0.26$	13.119	1.46
	126.	0.98105	350.	0.01895	$2.04 \pm 0.41$	24.494	1.09
	161.	0.98857	461.	0.01143	$1.43 \pm 0.24$	13.603	1.62
N7331	136.	0.99043	401.	0.00957	3	3.976	
	118.	0.99023	361.	0.00977	3	6.658	
N7552	144.	0.98818	418.	0.01182	$1.77 \pm 0.35$	20.019	1.43
	115.	0.99053	367.	0.00947	$2.21 \pm 0.46$	42.249	1.12
	147.	0.99277	483.	0.00723	$1.47 \pm 0.26$	21.235	1.67
N7793	176.	0.96455	381.	0.03545	$0.91 \pm 0.09$	0.869	2.29
	144.	0.95400	334.	0.04600	$1.02 \pm 0.20$	1.368	1.50
	181.	0.98096	440.	0.01904	$0.79 \pm 0.13$	0.900	2.62

<sup>a</sup>The parameters  $f_i$  are the mass fractions of the discrete components at the temperatures  $T_i$ . For each galaxy, the first line gives the results of the fits where  $T_1$  is constrained in order not to overestimate the column density, and the second (third) line where  $T_1$  is unrestricted. For galaxies where the S(2) flux is an upper limit, only constrained- $T_1$  fits were performed. When  $OPR_{\text{high T}} < 3$ ,  $T_2$  is fixed, and the free- $T_1$  fit results are then provided for two different values of  $T_2$ :  $1.14 \times T(\text{S1-S3})$  (second line) and  $1.5 \times T(\text{S1-S3})$  (third line). See Section 3.2 for explanations.

<sup>b</sup>The third component is characterized by  $T_3 = 1415$  K and  $f_3 = 0.00690$  (constrained- $T_1$  fit) or  $T_3 = 1455$  K and  $f_3 = 0.00477$  (free- $T_1$  fit).

<sup>c</sup> $T_3 = 1139$  K and  $f_3 = 0.00088$  (constrained- $T_1$  fit), or  $T_3 = 1319$  K and  $f_3 = 0.00019$  (free- $T_1$  fit).

<sup>d</sup> $T_3 = 1324$  K and  $f_3 = 0.00561$  (constrained- $T_1$  fit), or  $T_3 = 1414$  K and  $f_3 = 0.00010$  (free- $T_1$  fit).



Table 5. Masses in the warm and cold phases.<sup>a</sup>

galaxy	$M(T > T_1)$ ( $10^6 M_\odot$ )	$N_{\text{tot}}(\text{cold H}_2)$ ( $10^{20} \text{ mol. cm}^{-2}$ )	$M(T > T_1)$ / $M(\text{coldH}_2)$	CO beam (arcsec <sup>2</sup> )	corr.	map	ref CO
N24	3.13						
N337	11.61	9.4	0.18	1521.	1.52	IRAC	E96
N628	3.79	10.5	0.25	627.	1.04	BIMA	S05 (B93, Y95)
N855	0.50	2.3	0.22	415.	1.24	IRAC	W95
N925	1.15	< 18.7	> 0.06	386.			S05
N1097	103.01	183.	0.06	1590.	1.74	IRAC	Y95 (VV98, H93, K03)
N1266	12.97						
N1291	0.42	11.7	0.04	1452.	2.33	IRAC	T91
N1316	0.85	19.9	0.006	1452.	2.55	IRAC	H01
N1482	73.91	136.	0.04	1521.	1.76	IRAC	E96
N1512	6.34						
N1566	10.22						
N1705	0.50						
N2403	0.33	23.8	0.10	855.	0.96	BIMA	E96
N2798	34.54	46.5	0.11	855.	2.50	IRAC	E96
N2841	0.74	< 27.4	> 0.03	370.			S05
N2915	0.0077	< 6.0	> 0.02	1452.	2.04	IRAC	A04
N2976	0.28	10.6	0.20	452.	1.12	BIMA	A04 (I05, Y95)
N3031	0.10	< 7.9	> 0.09	3068.			S05
N3049	3.39	31.3	0.02	573.	1.48	IRAC	C97
N3184	2.09	40.1	0.07	441.	1.23	BIMA	S05 (Y95, S93)
N3190	11.33	37.8	0.09	380.	1.14	IRAC	L98
N3198	3.54	33.1	0.10	415.	1.31	IRAC	B93
N3265	6.45	25.8	0.06	227.	0.76	IRAC	G91
Mrk33 <sup>b</sup>	21.35	16.5	0.23	380.	1.15	IRAC	S92 (A04)
N3351	11.30	95.1	0.05	606.	0.85	BIMA	P97 (S05)

Table 5—Continued

galaxy	$M(T > T_1)$ ( $10^6 M_\odot$ )	$N_{\text{tot}}(\text{cold H}_2)$ ( $10^{20} \text{ mol. cm}^{-2}$ )	$M(T > T_1)$ / $M(\text{coldH}_2)$	CO beam (arcsec <sup>2</sup> )	corr.	map	ref CO
N3521	4.95	26.1	0.22	201.	0.68	BIMA	N01 (Y95)
N3621	1.88						
N3627	2.39	196.	0.01	271.	0.95	BIMA	S05 (Y95)
N3773	2.98	< 13.5	> 0.12	2376.	5.83	IRAC	L05
N3938	4.23	26.4	0.10	262.	0.96	BIMA	S05 (Y95)
N4125	1.20	6.5	0.04	415.	1.13	IRAC	W95
N4254	7.86	81.8	0.02	201.	0.75	S99	N01 (K88)
N4321	47.83	133.	0.03	415.	0.74	BIMA	B93 (S05, K96, H93)
N4450	5.04	16.4	0.07	131.	0.90	BIMA	S05
N4536	102.41	50.2	0.11	855.	1.08	S03	E96
N4552	1.91						
N4559 <sup>c</sup>	9.29	13.9	0.45	355.	0.98	BIMA	S05
N4569	17.48	200.	0.02	118.	0.73	BIMA	S05 (N01, K88, H93)
N4579	2.36	36.7	0.01	118.	0.42	BIMA	S05 (K88)
N4594	0.47	< 6.2	> 0.04	346.			B91
N4625	1.85	9.2	0.20	452.	1.10	IRAC	A04 (B03)
N4631	8.44	96.2	0.10	1590.	1.76	IRAC	Y95
N4725	5.85	< 54.8	> 0.03	210.			S05
N4736	1.20	81.3	0.05	767.	1.37	BIMA	S05 (S98, Y95)
N4826	3.00	250.	0.04	1337.	1.81	BIMA	S05 (N01, Y95)
N5033	10.86	96.0	0.05	855.	1.39	BIMA	C01 (S05, N01, E96, A95)
N5055	5.98	122.	0.06	645.	1.23	BIMA	S05 (M99, E96, A95, Y95, S93)
N5194	1.46	40.0	0.04	474.	0.98	BIMA	S05
N5194_A	0.95						
N5194_B	3.37						
N5194_C	6.07						

Table 5—Continued

galaxy	$M(T > T_1)$ ( $10^6 M_\odot$ )	$N_{\text{tot}}(\text{cold H}_2)$ ( $10^{20} \text{ mol. cm}^{-2}$ )	$M(T > T_1)$ / $M(\text{coldH}_2)$	CO beam (arcsec <sup>2</sup> )	corr.	map	ref CO
N5194_D	3.53						
N5194_E	3.33						
N5194_F	3.07						
N5194_G	2.91						
N5195	5.51	110.	0.06	177.	0.58	IRAC	K02 (S89)
N5713	30.27	50.8	0.08	1590.	2.01	IRAC	Y95 (Y03)
N5866	2.80	50.9	0.03	346.	1.02	IRAC	W03 (T94)
N6822_A <sup>b</sup>	0.0047	7.8	0.14	415.	1.26	IRAC	I03
N6822_B	0.0062						
N6822_C	0.0030						
N6946	11.95	486.	0.03	816.	0.95	BIMA	S05
N7331	9.69	56.0	0.07	346.	0.94	BIMA	I99 (N01)
N7552	269.01	171.	0.12	1452.	1.86	IRAC	A95 (C92)
N7793	0.10	25.4	0.03	1452.	1.87	IRAC	I95

<sup>a</sup>We retain here results from constrained- $T_1$  fits, giving lower masses in the warm phase. Whenever the excitation diagrams are ambiguous regarding the value of  $OPR_{\text{high T}}$ , we adopt the mass of warm  $\text{H}_2$  derived with  $OPR_{\text{high T}} < 3$ , which is also smaller than the mass derived with  $OPR_{\text{high T}} = 3$  (see Table 4). References for CO intensities are given in abbreviated form. The quantity "corr." designates the factor that was applied to the CO brightness in order to correct for the different apertures of the CO and  $\text{H}_2$  observations, and "map" is the image used to derive this factor (see Section 4).

<sup>b</sup>Mrk 33: The  $\text{H}_2$  mass fraction in the warm phase may be severely overestimated in this blue compact galaxy, because of the use of an inappropriate factor to convert CO flux to total  $\text{H}_2$  mass. Israel (2005) derive a factor about four times higher than the standard factor used here, which would make the total  $\text{H}_2$

mass also four times higher. NGC 6822-A: In their detailed study of this region, Israel et al. (2003) derive a CO flux to total H<sub>2</sub> mass conversion factor twenty times higher than the standard factor, so that the same remark as for Mrk 33 applies.

<sup>c</sup>NGC 4559 has the highest mass fraction in the warm phase, even when constraining  $T_1$  to be near-maximal. The warm phase is dominated by gas at  $\sim 100$  K. NGC 4559 is a quiescent galaxy, as obvious from the low  $F_{24}/F_{7.9}$  flux density ratio within the IRS aperture. The galaxy has little molecular gas, and it is possible that it is mostly diffuse, and thus more exposed to FUV radiation. The same remarks may apply, to a lesser degree, to other galaxies of the sample such as NGC 628.

Table 6. Nuclear X-ray fluxes in the 2-10 keV band and references.

galaxy	$F_X$ ( $10^{-17} \text{ W m}^{-2}$ )	ref	note <sup>a</sup>
N628	0.59	K05	2
N1097	173.	N06	1
N1291	7.45	K05	2
N1316	10.1	KF03	2
N2841	1.06	HF01	1
N3031	1020.	HF01	1
N3184	1.12	K05	2
N3627	< 0.32	HF01	1
N4125	1.07	S04	1
N4321	< 1.25	HF01	1
N4552	12.1	SD05	1
N4569	7.61	HF01	1
N4579	264.	HF01	1
N4594	120.	HF01	1
N4725	7.15	HF01	1
N4736	27.4	SD05	1
N4826	< 1.08	HF01	1
N5033	123.	HF01	1
N5194	1300.	S04	1
N5195	4.45	K05	2
N5866	0.18	SD05	1
N7331	2.28	S04	1

<sup>a</sup>1: Quantity given directly in the cited paper (abbreviated notation as in the list of references).  
2: Quantity computed as in Ho et al. (2001) (HF01) from counts in the 0.2-8 keV band.

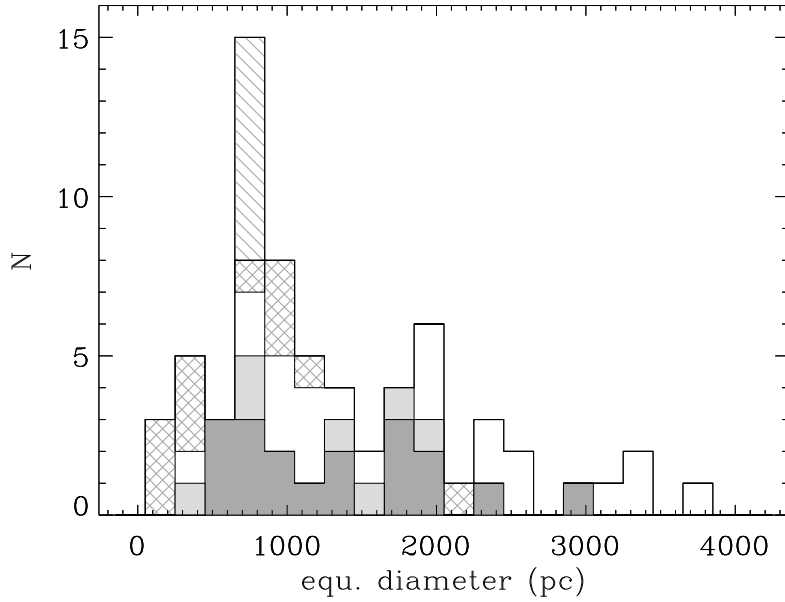


Fig. 1.— Histogram of the equivalent diameters of the areas over which the line and continuum fluxes were integrated. The apertures were defined by the intersection of the spectral maps in the LH, SH and SL modules of the IRS instrument. The extranuclear regions in NGC 5194 are shown by the hatched histogram, the regions within dwarfs by the cross-hatched histogram, the LINER nuclei by the darker shade and the Sy nuclei by the lighter shade.

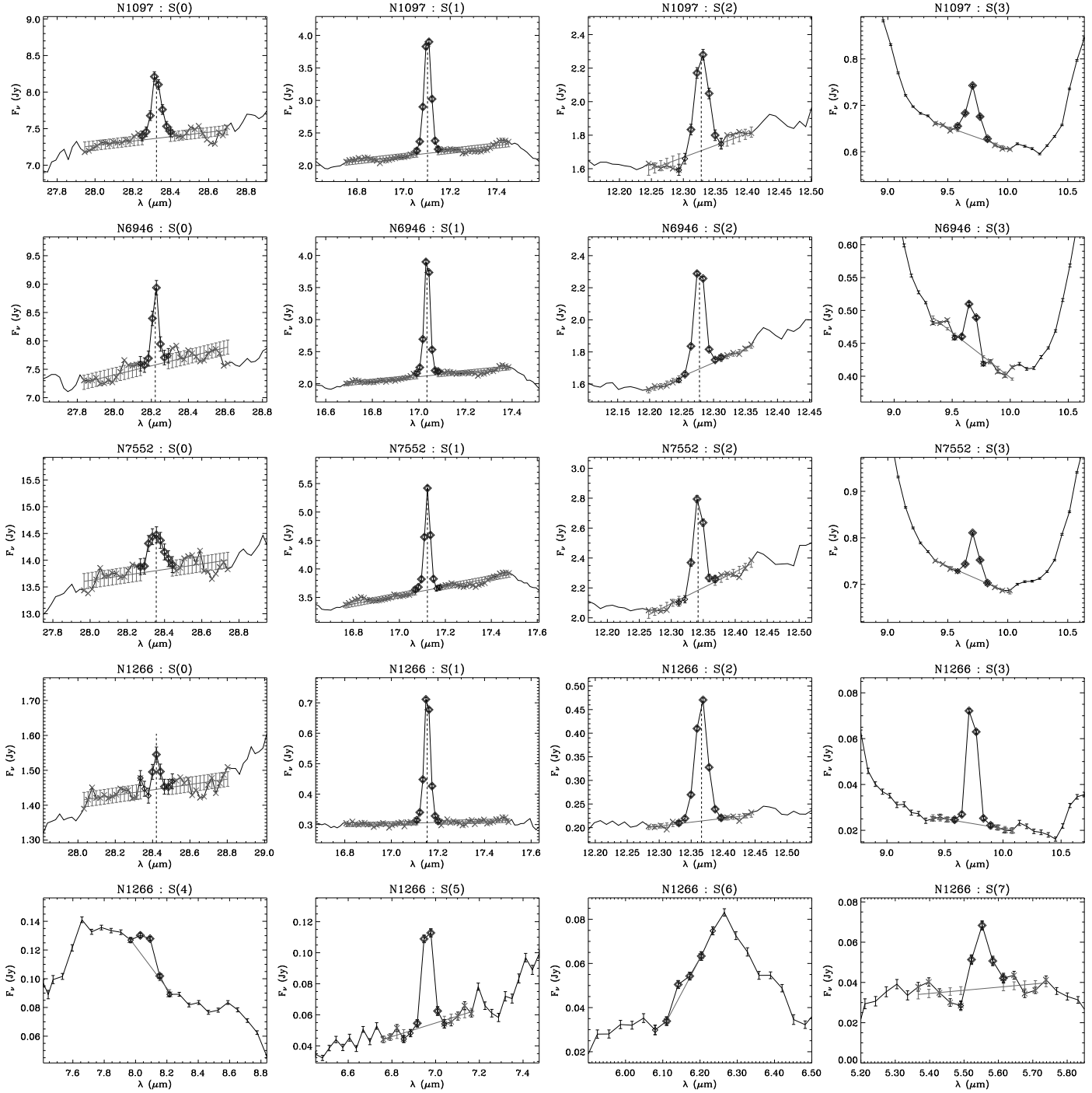


Fig. 2.— Examples of H<sub>2</sub> line spectra: the circumnuclear starbursts NGC 1097, NGC 6946 and NGC 7552, and the three galaxies for which we could estimate the fluxes of higher transitions than S(3), NGC 1266, NGC 4569 and NGC 4579. The straight line indicates the fitted pseudo-continuum, and the diamond symbols show the wavelength range over which the line flux was integrated.

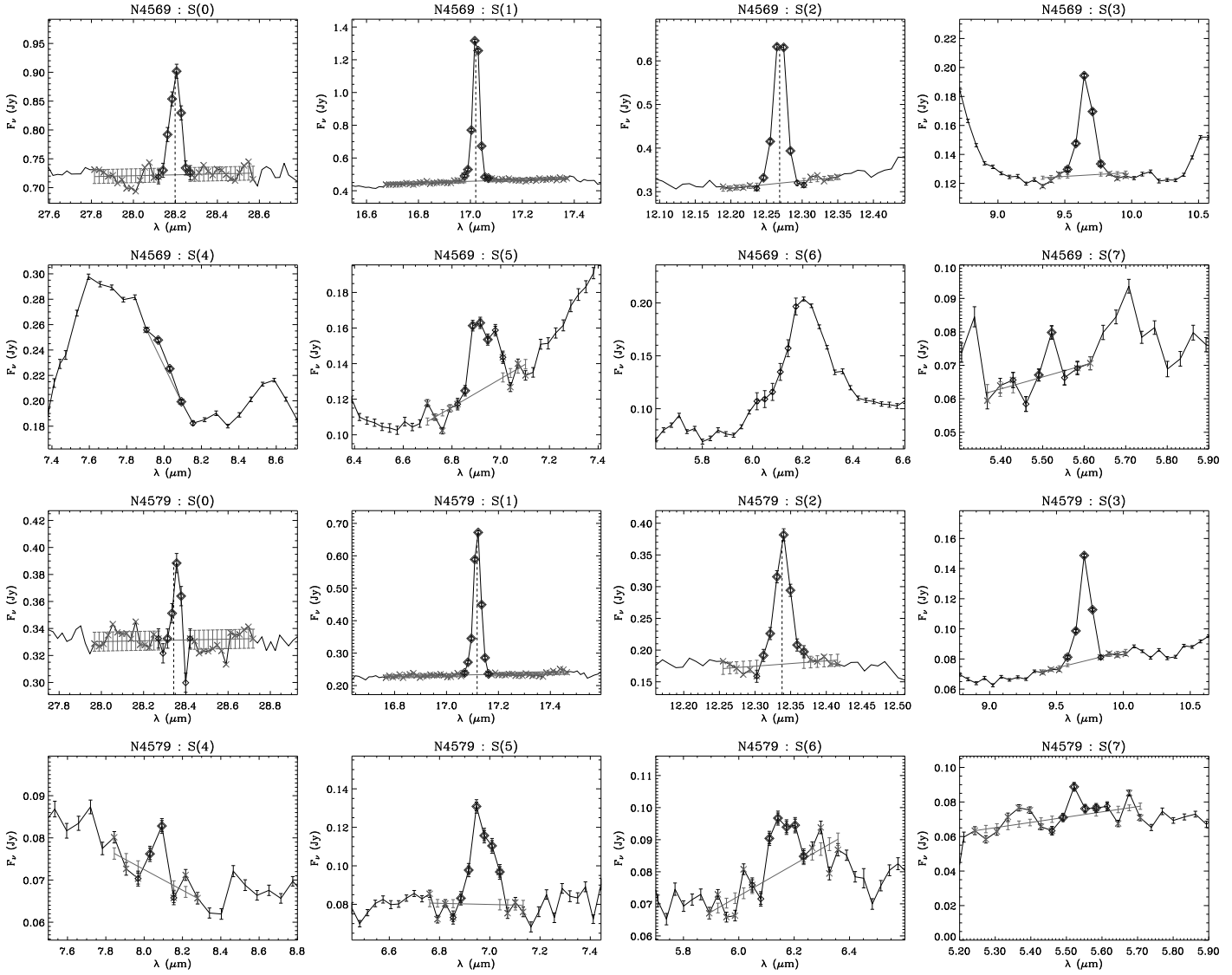


Fig. 2.— (continued).



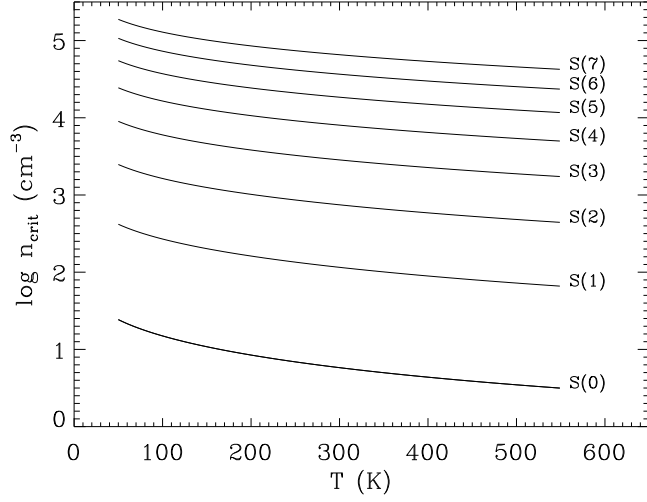


Fig. 3.— Critical densities for collisional deexcitation by  $\text{H}_2$  for the rotational transitions in the wavelength range of the IRS instrument.

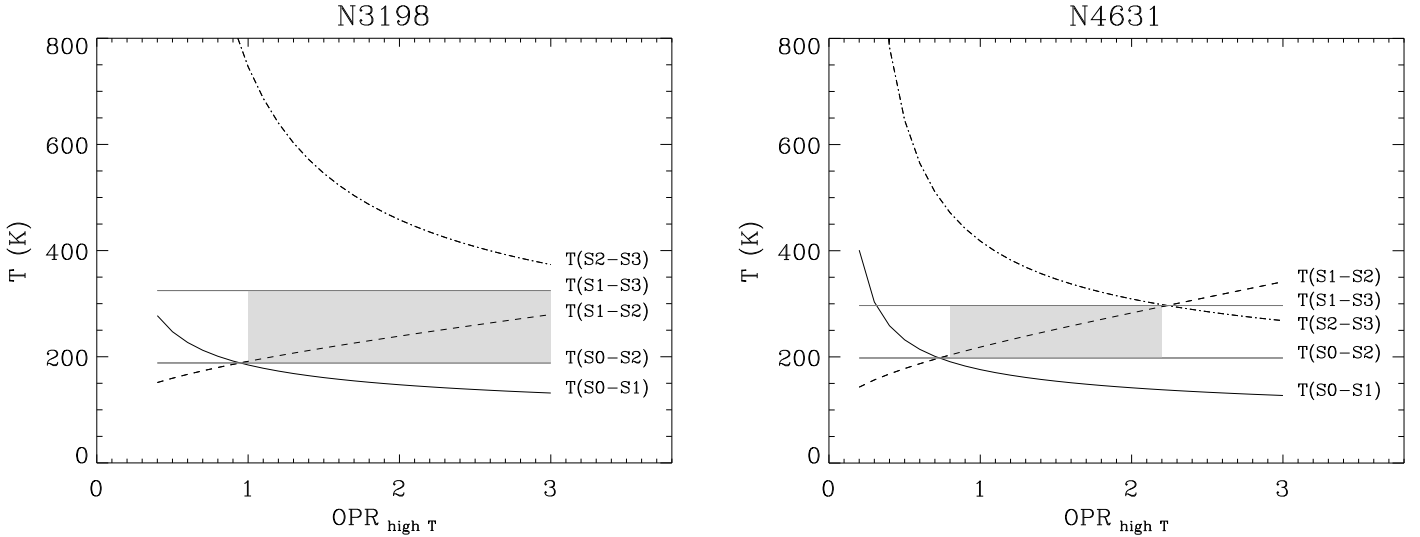


Fig. 4.— Examples of how the possible range of apparent  $OPR_{\text{high T}}$  is determined for each galaxy. Thermalization requires the apparent temperatures derived from each pair of transitions to be monotonic as a function of upper level energy; in particular, the conditions  $T(S1 - S2) < T(S1 - S3) < T(S2 - S3)$  (indicated by the shaded areas) have to be satisfied. The apparent temperatures derived for NGC 3198 are compatible with  $OPR_{\text{high T}} = 3$ , whereas they require  $OPR_{\text{high T}} < 3$  for NGC 4631 (see text).

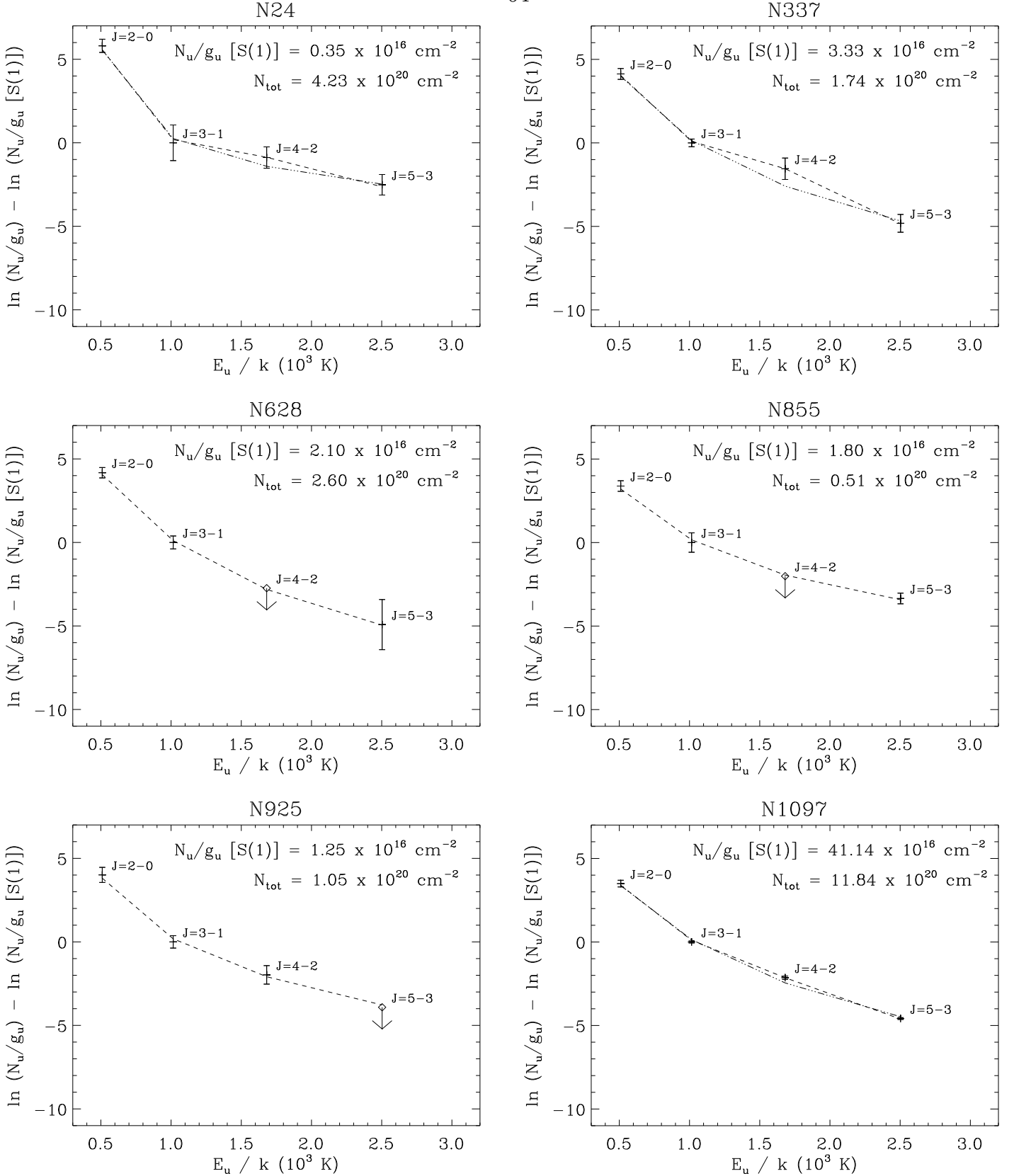


Fig. 5.— Excitation diagrams. The  $N_u/g_u$  ratios are normalized by the S(1) transition. The dashed line indicates the best fit (see text). Whenever  $OPR_{\text{highT}} < 3$  is required by the temperature constraints, a dot-dash line shows for comparison the fit obtained with  $OPR_{\text{highT}} = 3$ .

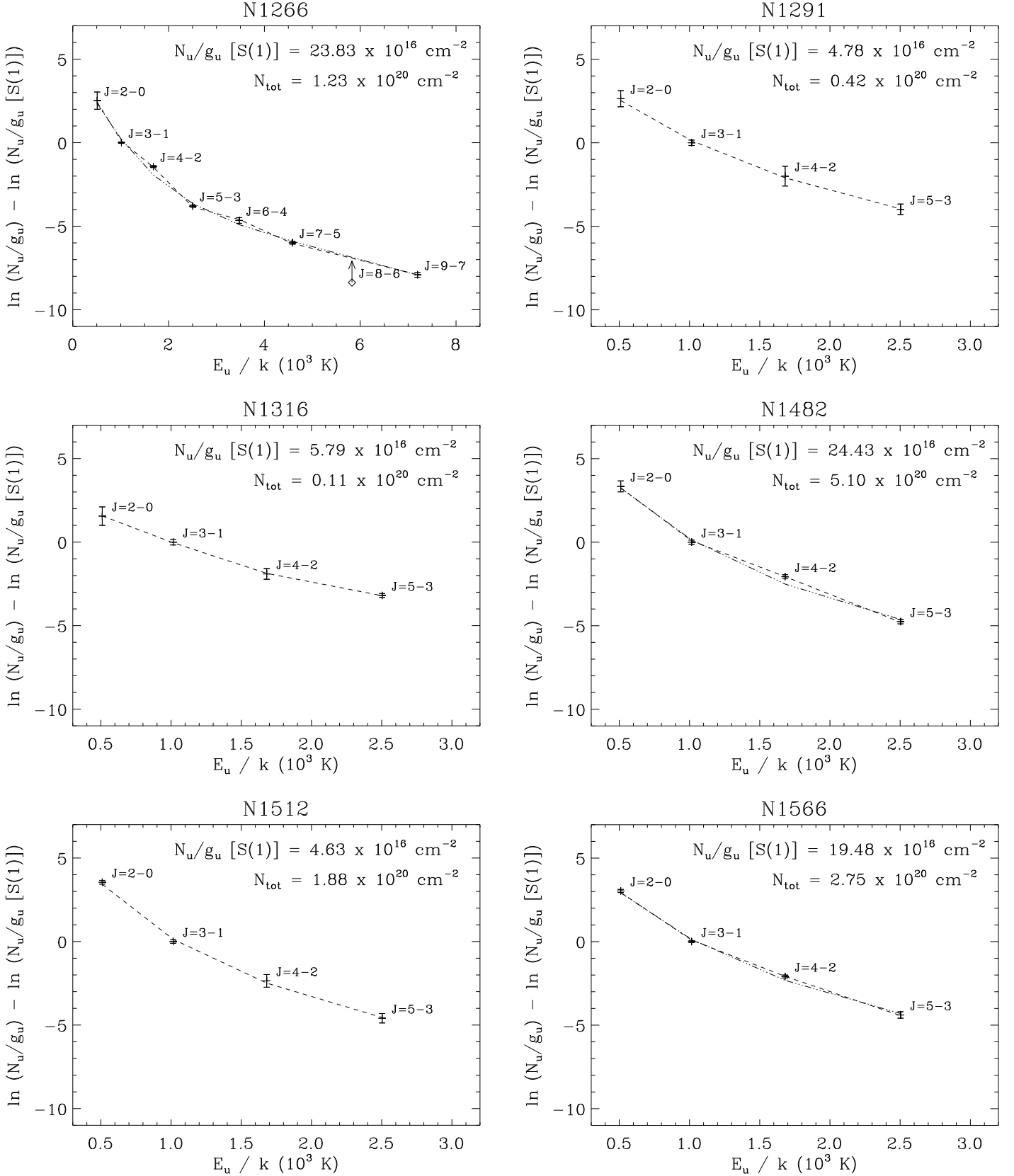


Fig. 5.— (continued).

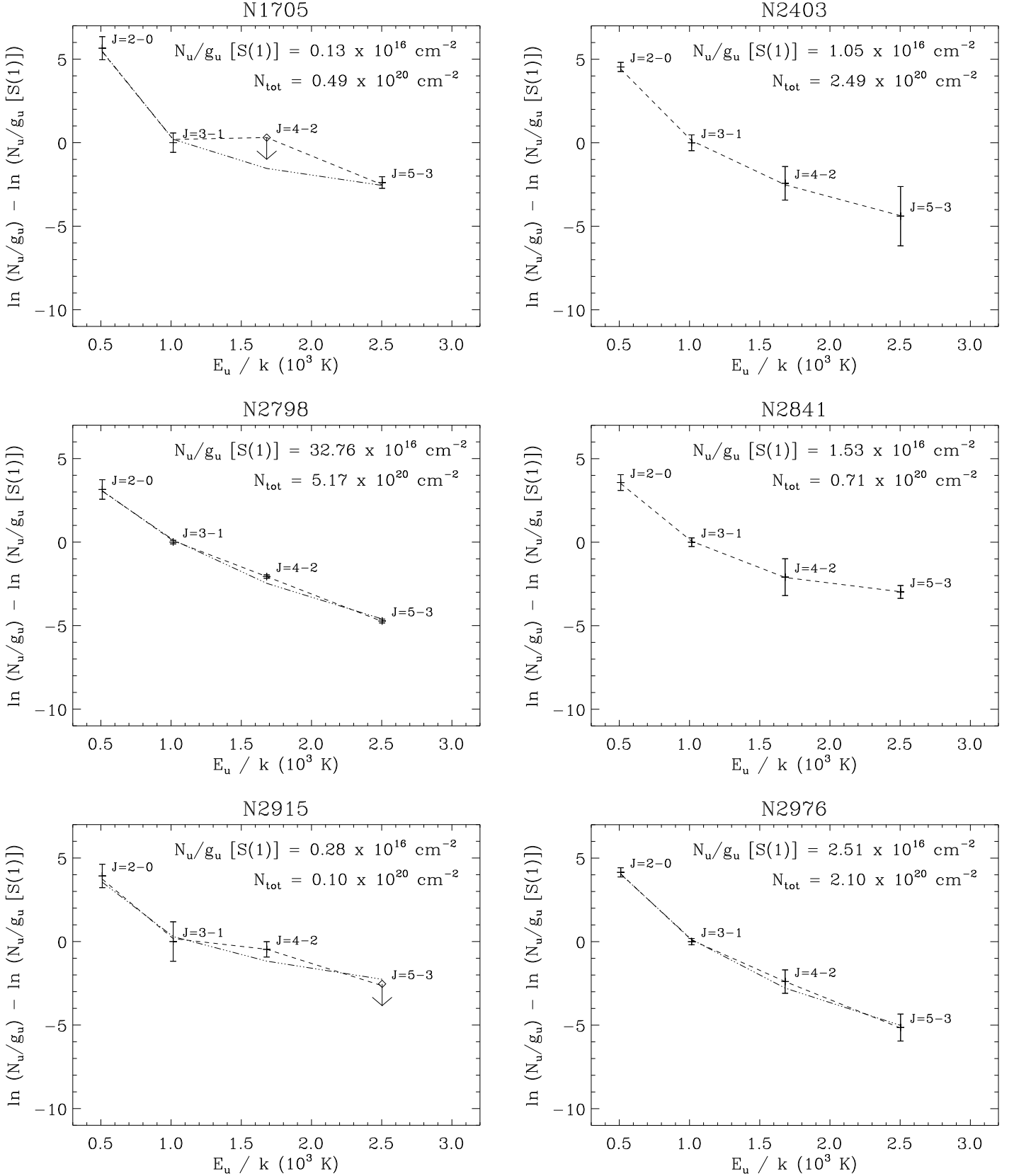


Fig. 5.— (continued).

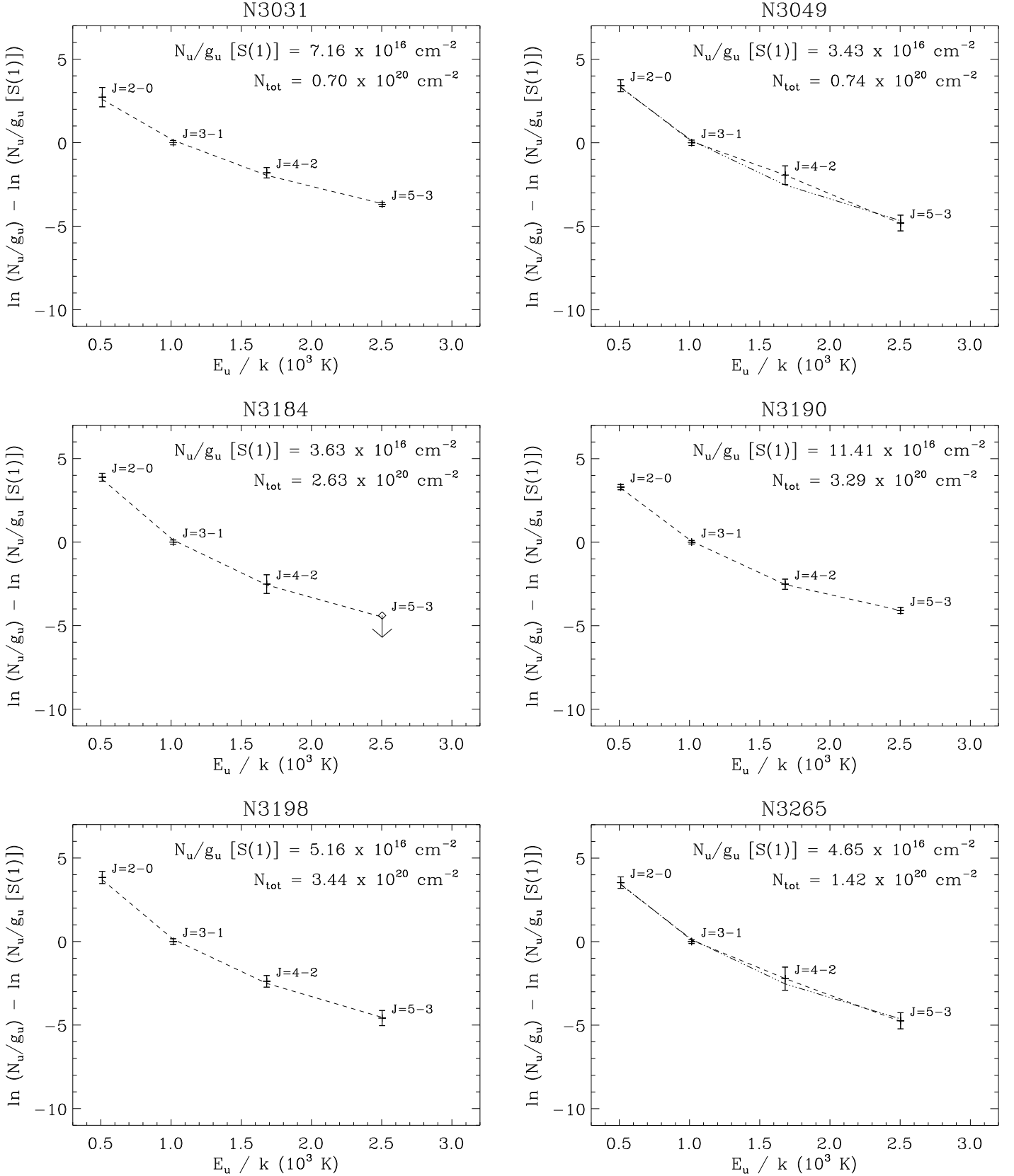


Fig. 5.— (continued).

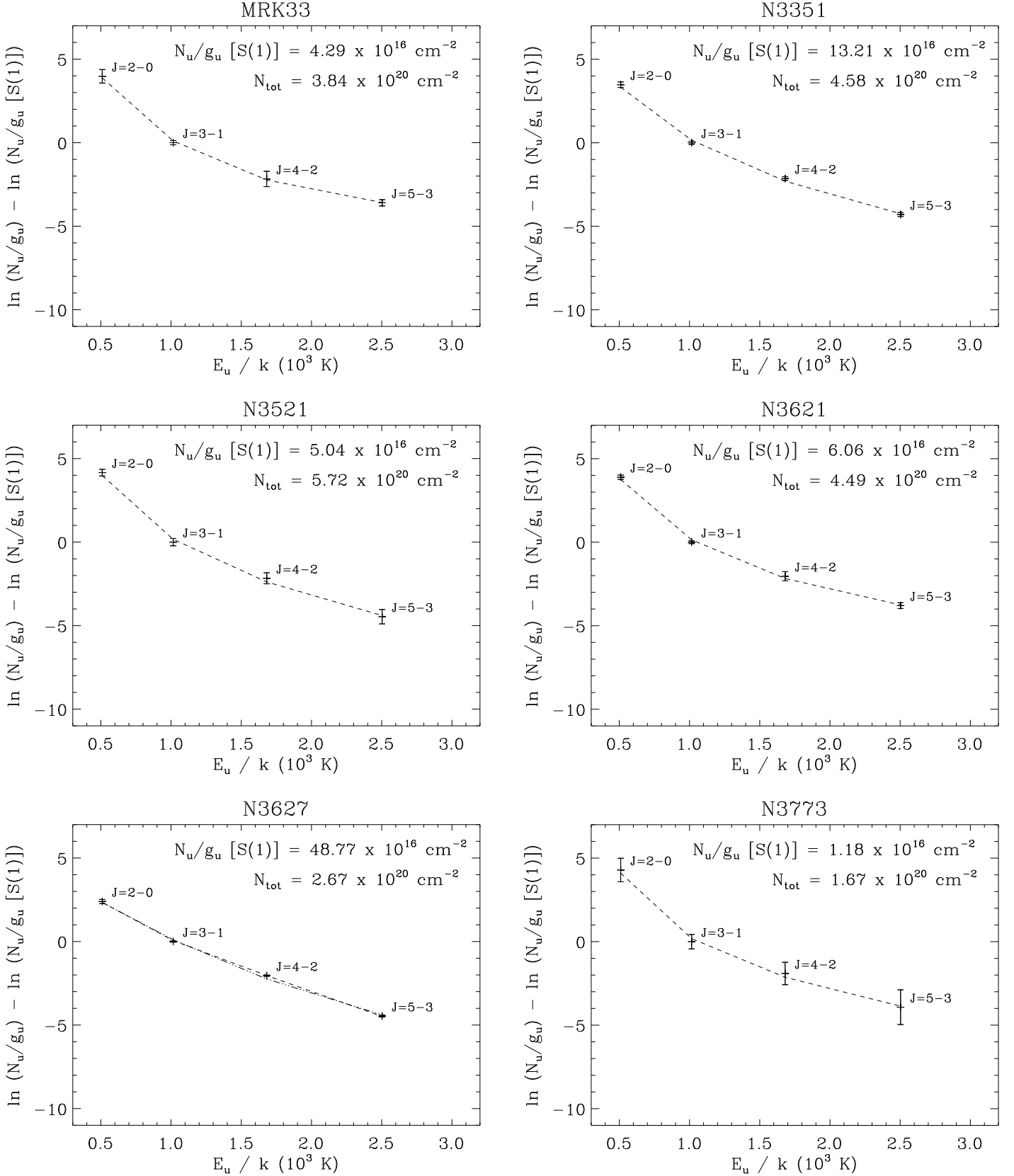


Fig. 5.— (continued).

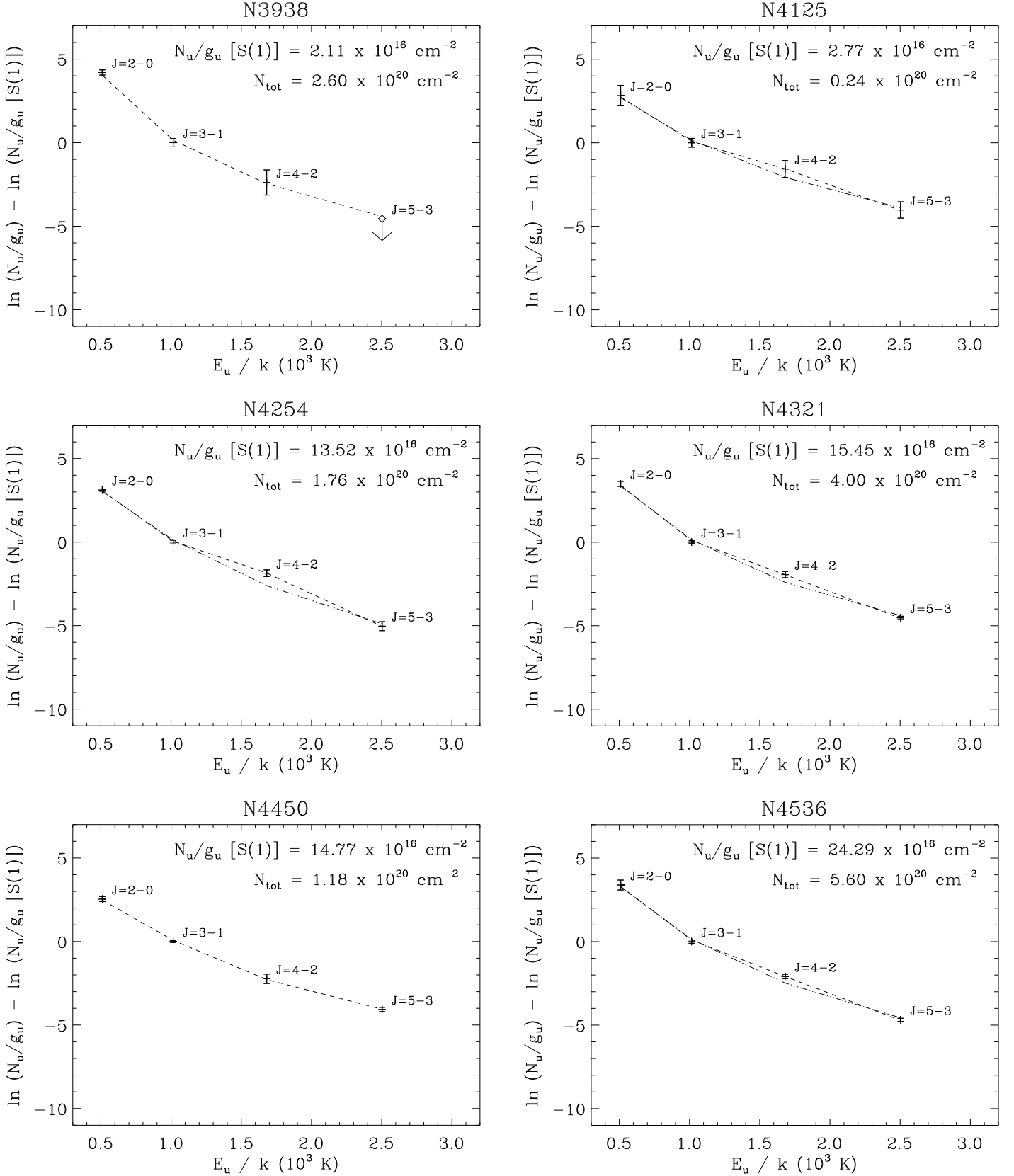


Fig. 5.— (continued).

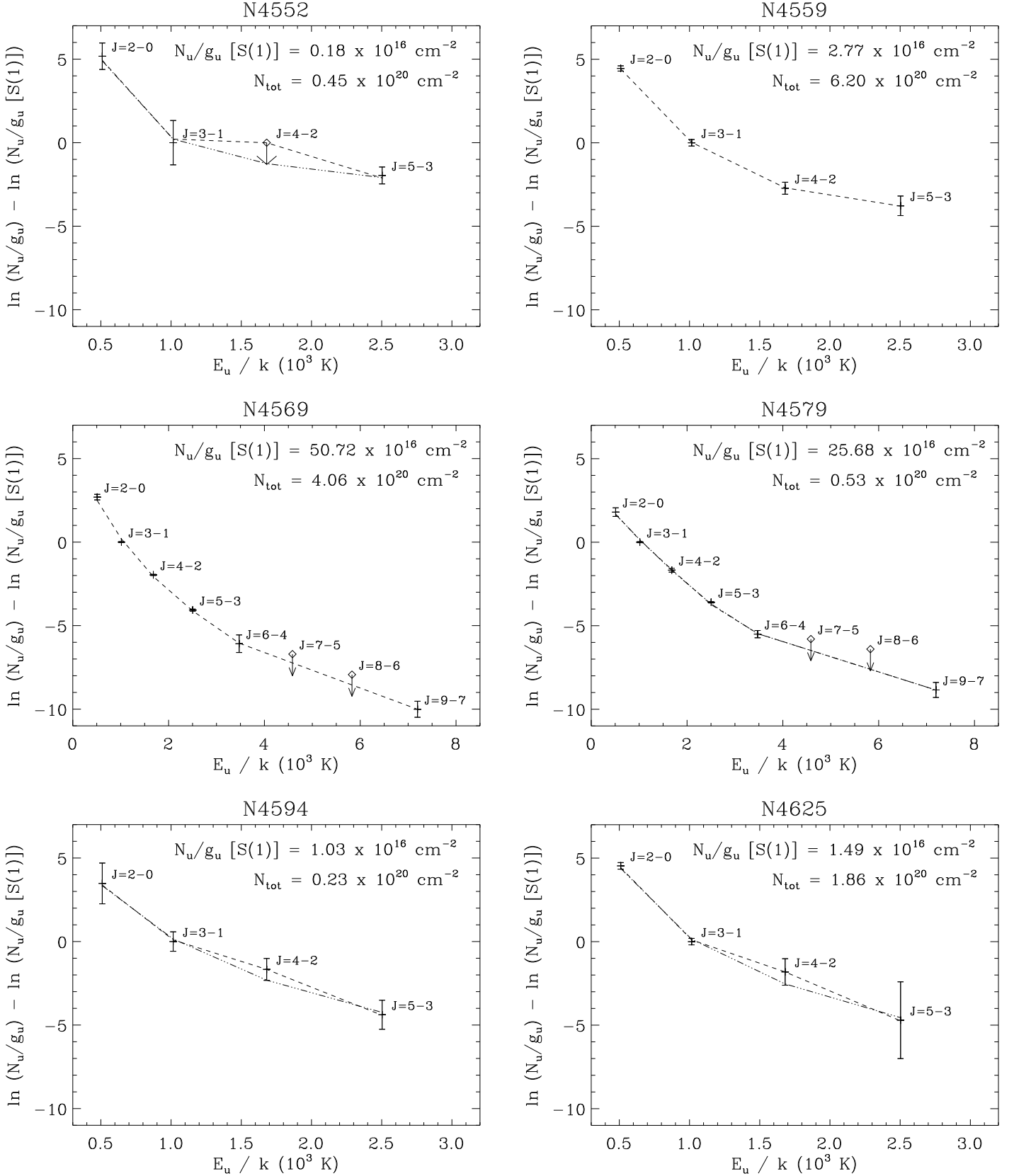


Fig. 5.— (continued).



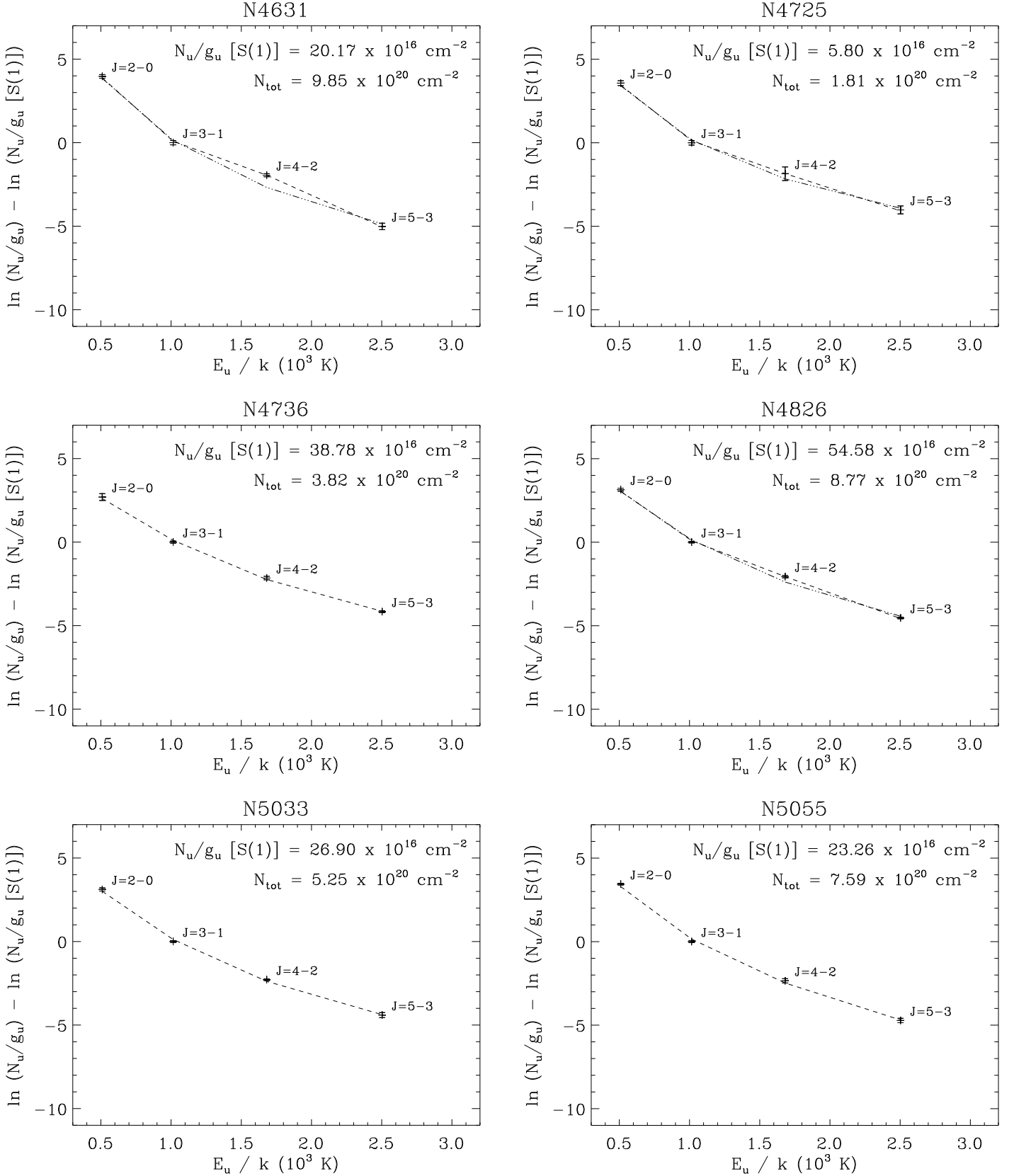


Fig. 5.— (continued).

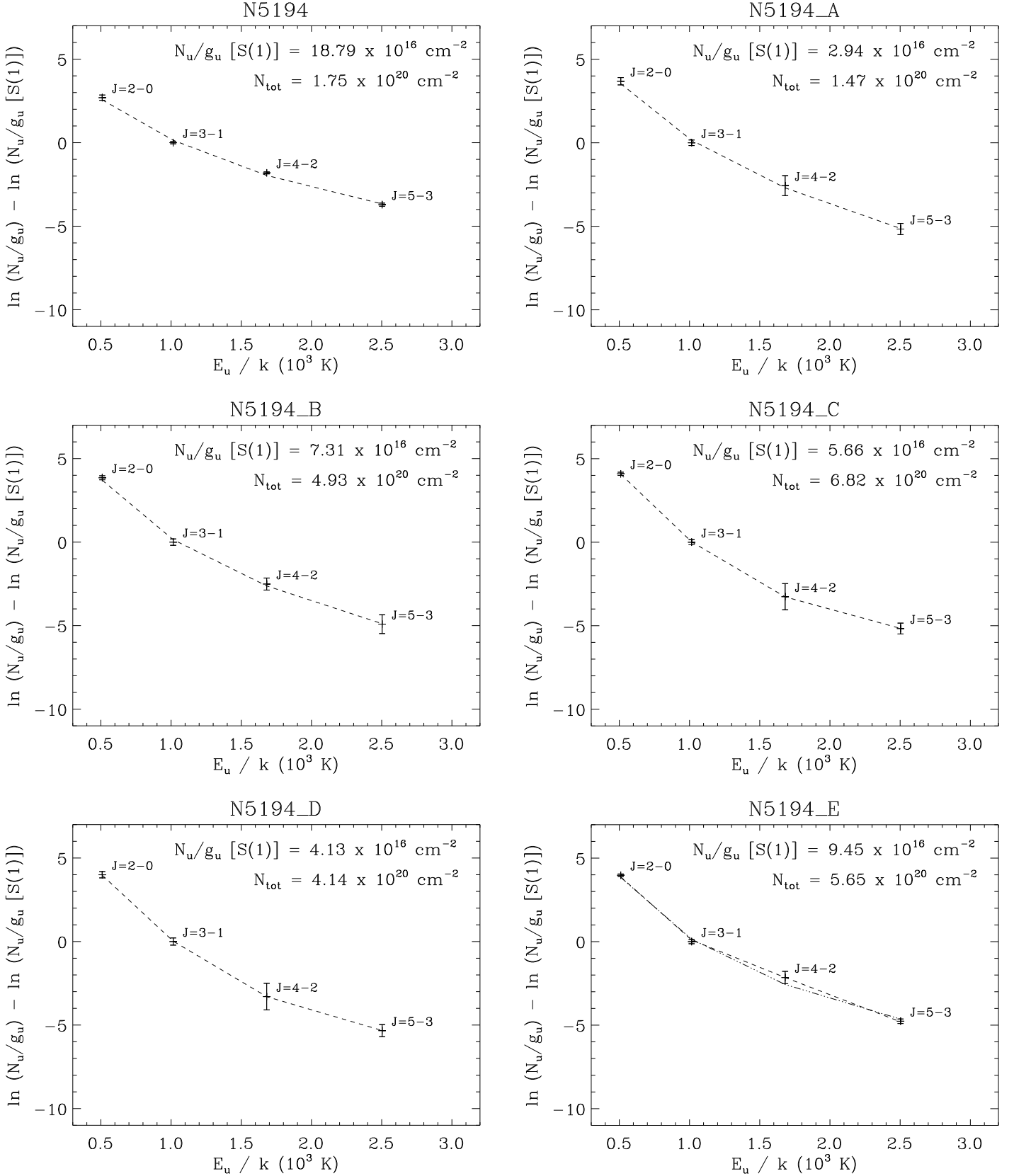


Fig. 5.— (continued).

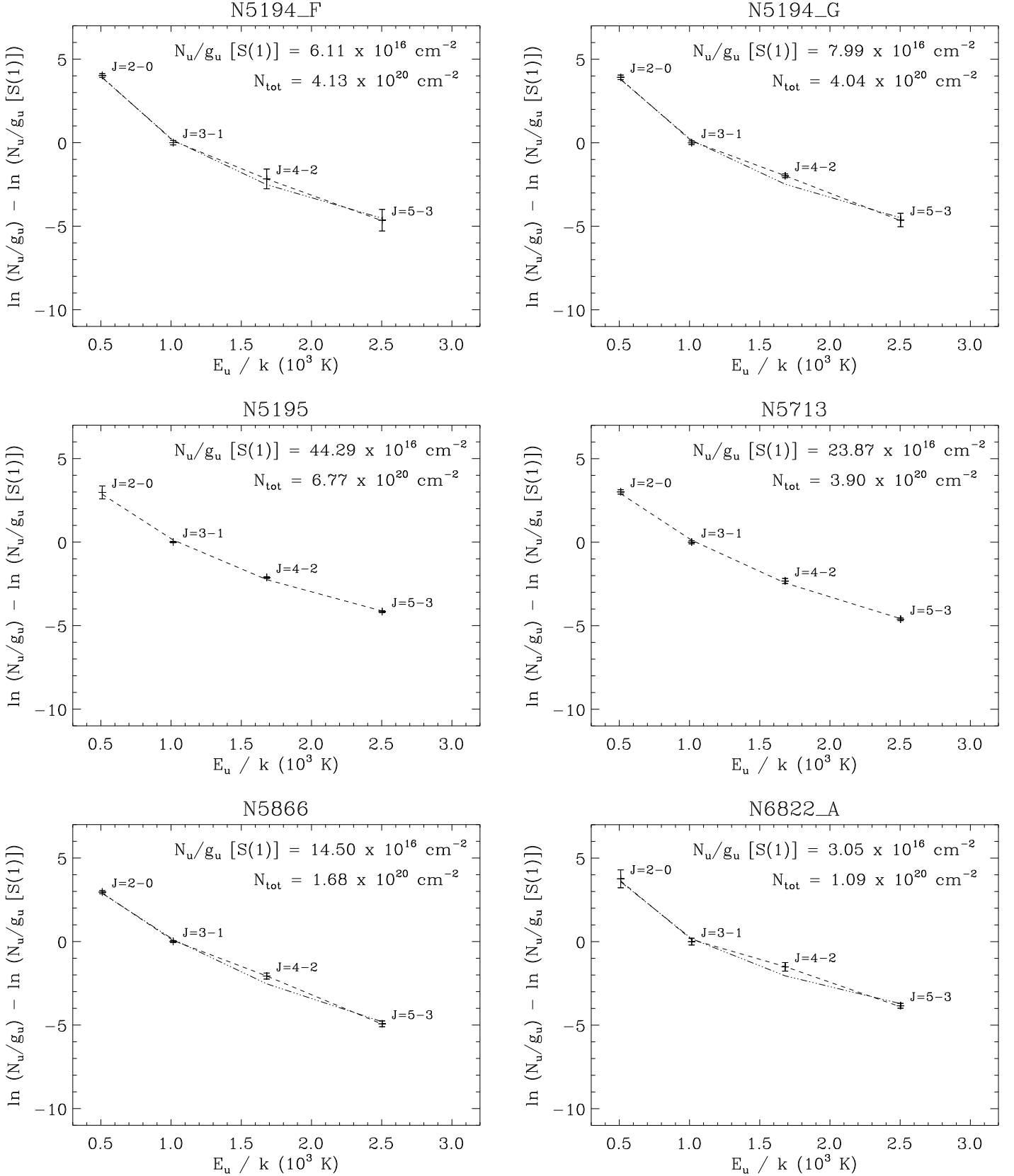


Fig. 5.— (continued).

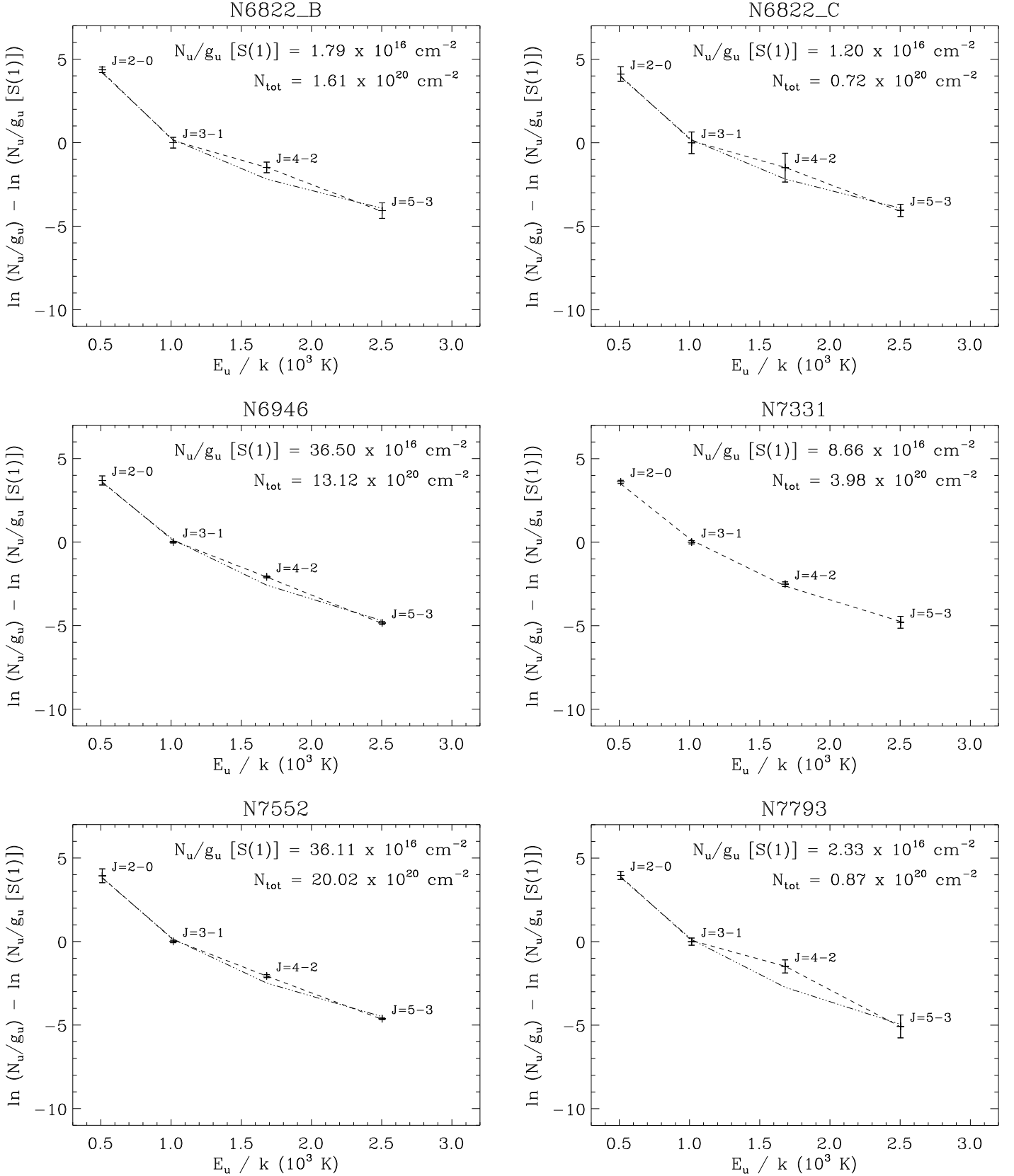


Fig. 5.— (continued).

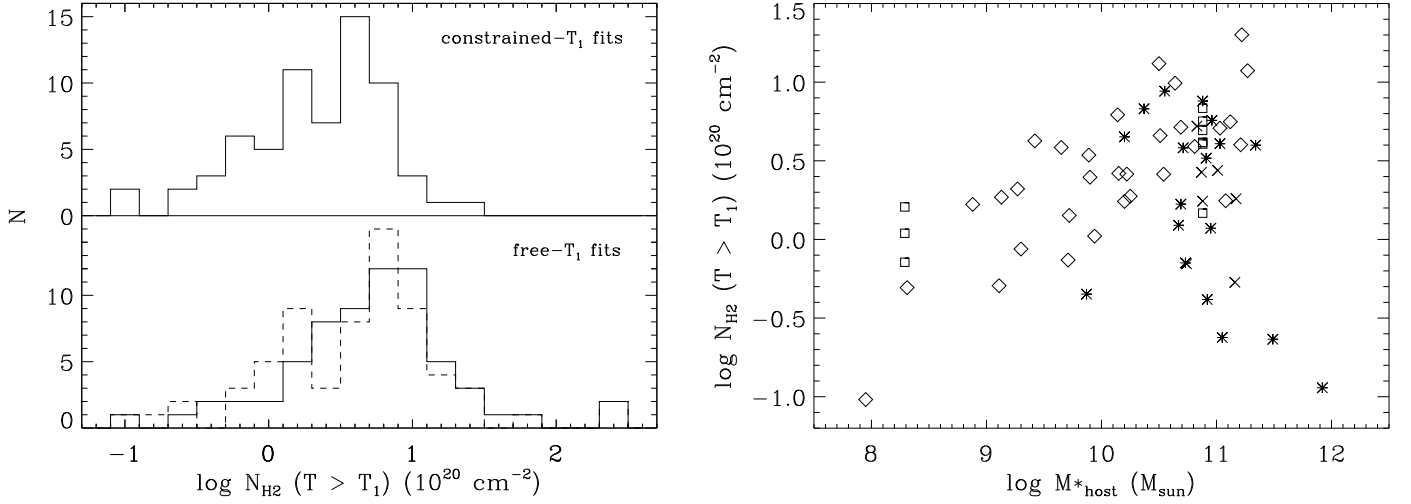


Fig. 6.— **Left:** Histograms of the derived column densities in fits where the lower temperature is constrained (top) and where it is free to vary (bottom). In the latter case, the solid line indicates results with low  $T_2$  values and the dashed line with high  $T_2$  values, in  $OPR_{\text{highT}} < 3$  fits (see text). **Right:** Column densities as a function of the total stellar mass of the host galaxy, estimated as by Lee et al. (2006). Star-forming nuclei are represented as diamonds, the extranuclear regions in NGC 5194 and NGC 6822 as squares, LINER nuclei as star symbols and Sy nuclei as crosses. By definition, dwarf galaxies have  $M_{\text{host}}^* < 10^{9.7} M_{\odot}$ .

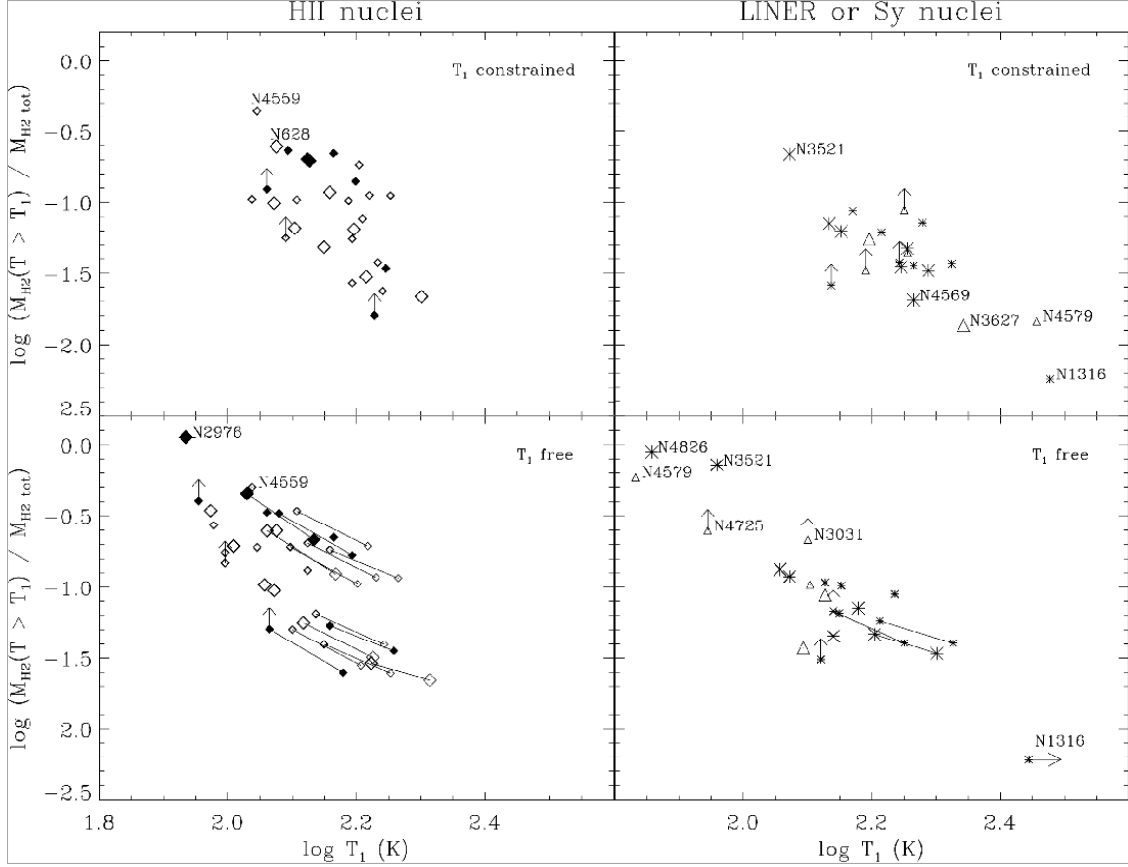


Fig. 7.— Fraction of H<sub>2</sub> in the warm phase ( $T > T_1$ ) as a function of  $T_1$ , the lower temperature of the two components fitted to the rotational lines. Nuclei classified as purely star-forming are shown as diamonds, and regions in dwarf galaxies as filled diamonds; nuclei classified as LINER or Sy are shown as star symbols or triangles, respectively. The symbol size is increased for targets with the most robust CO flux estimates (see text). In the free- $T_1$  fits with  $OPR_{\text{high } T} < 3$ , results with the two values of  $T_2$  considered here (see text) are connected by a line segment; the fits with  $T_2 = 1.5 \times T(S1 - S3)$  produce higher temperatures and lower warm H<sub>2</sub> masses than the fits with  $T_2 = 1.14 \times T(S1 - S3)$ .

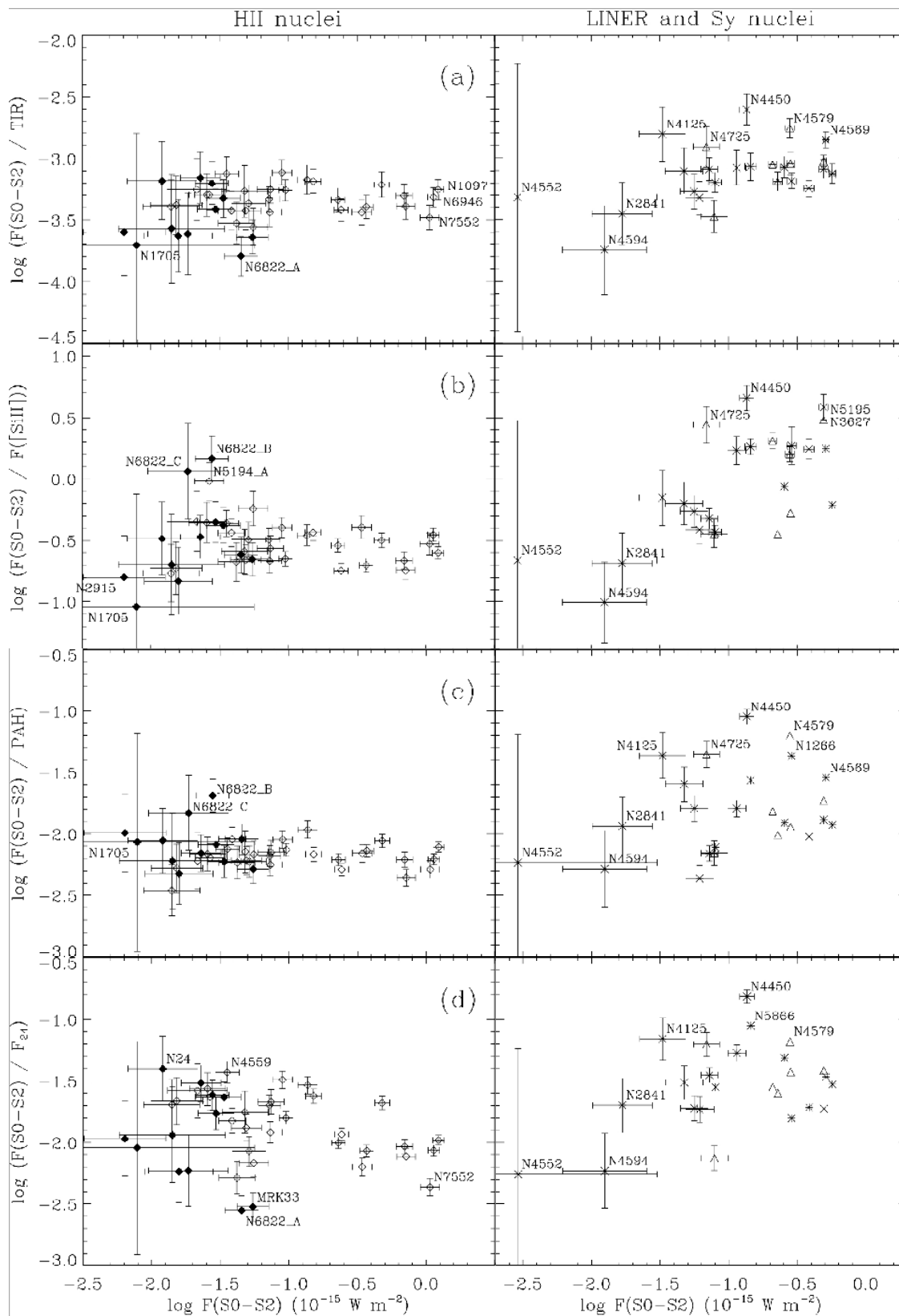


Fig. 8.— Ratio of the power emitted in the sum of the S(0) to S(2) transitions to: (a) the total infrared power; (b) the [SII] line power; (c) the power emitted in the aromatic bands within the IRAC4 filter (assuming a filter width of 13.9 THz and subtracting the stellar emission as mentioned in Sect. 2.2); (d) the power emitted at 24  $\mu$ m within the MIPS1 filter (assuming a filter width of 3.1 THz). The symbol coding is as in Fig. 7. Excluding the regions within NGC 6822 (the targets with the smallest projected aperture), as well as NGC 1705 and NGC 2915 (the galaxies with the smallest H<sub>2</sub> brightness), the average and dispersion of each logarithmic power ratio is: (a)  $-3.36 \pm 0.15$  for HII nuclei and complexes; and  $-3.12 \pm 0.24$  for LINER and Sy nuclei; (b)  $-0.53 \pm 0.17$  and  $-0.04 \pm 0.42$ , respectively; (c)  $-2.19 \pm 0.10$  and  $-1.80 \pm 0.34$ ; (d)  $-1.85 \pm 0.28$  and  $-1.51 \pm 0.31$ . Note that the H<sub>2</sub> flux (abscissa) is almost proportional to the H<sub>2</sub> brightness, because of the quasi-uniform beam.

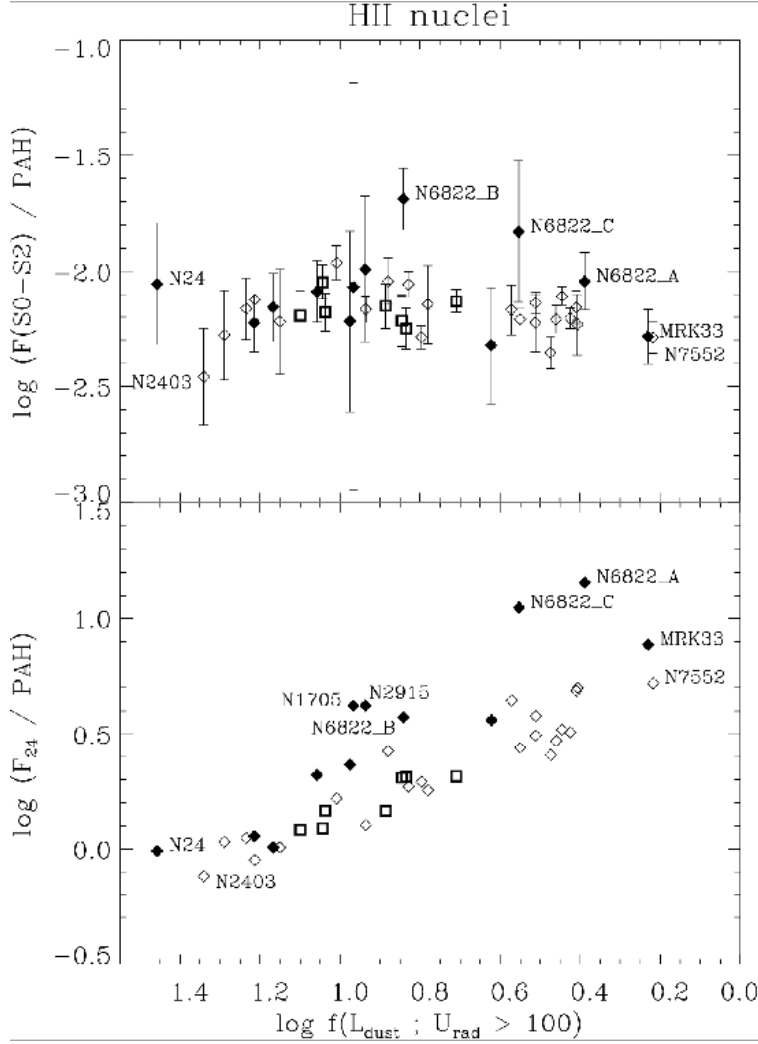


Fig. 9.— Power ratios of  $H_2$  to aromatic bands (shown in Fig. 8c) and  $24\mu m$  emission to aromatic bands as a function of  $P_{24} = 1.05((\nu_{24}F_{24} - 0.14\nu_{7.9}F_{7.9\text{dust}})/(\nu_{71}F_{71} + \nu_{156}F_{156}) - 0.035)^{0.75}$ , a quantity closely related to  $f(L_{dust} ; U_{rad} > 100)$  defined by Draine et al. (2007), which is the fraction of the total dust luminosity coming from regions with radiation field intensities more than 100 times the average local field. Only HII nuclei and complexes are shown. Dwarfs are represented by filled diamonds and the extranuclear regions within NGC 5194 by thick squares.



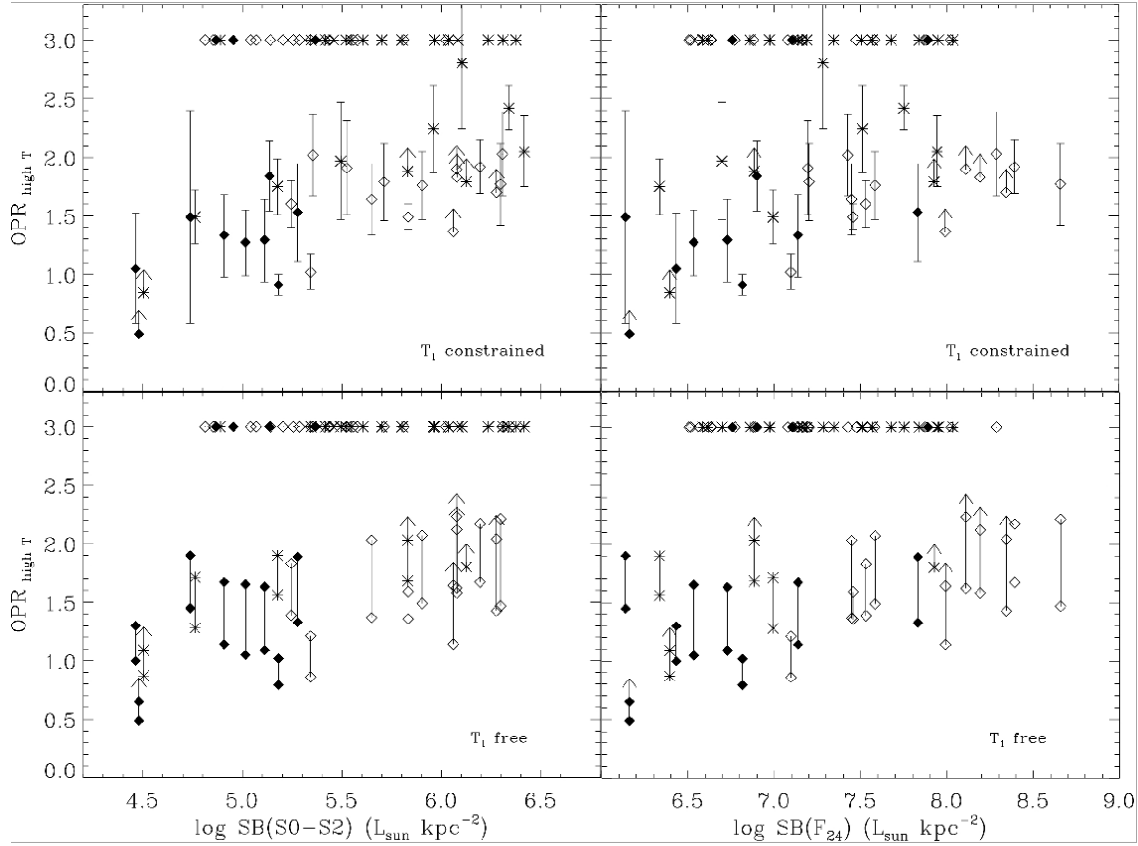


Fig. 10.— Ortho to para ratios as a function of the surface brightness in the sum of the S(0) to S(2) transitions and as a function of the surface brightness in the  $24\mu\text{m}$  band. For a number of sources, based on the excitation diagram, we adopted a fixed  $OPR_{\text{high } T} = 3$  (see text). Galaxies with upper limits in the S(2) line, or with indications of possibly non-negligible optical depth at  $10\mu\text{m}$ , are shown as lower limits of  $OPR_{\text{high } T}$ . For free- $T_1$  fits, results obtained from the two values of  $T_2$  considered here are connected by a line segment. The symbol coding is as in Fig. 7, except that LINER and Sy nuclei are both shown as star symbols.

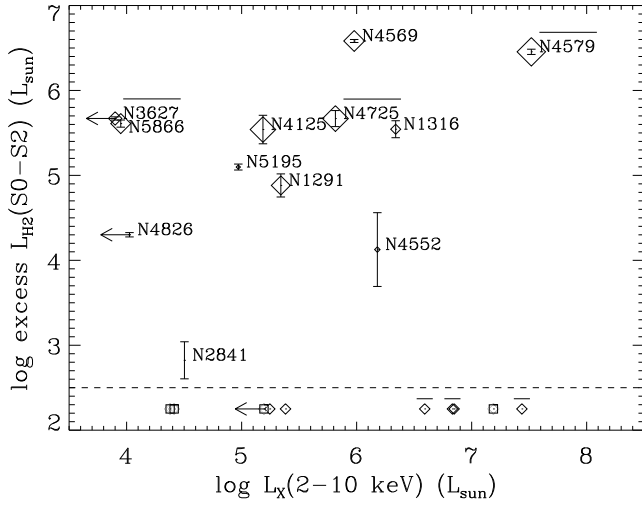


Fig. 11.— Excess H<sub>2</sub> luminosity as a function of 2-10 keV X-ray luminosity. The excess H<sub>2</sub> emission is defined from the relation between H<sub>2</sub> and aromatic band power, shown in Figure 8c, as the difference between the total H<sub>2</sub> emission and the quantity  $10^{-1.94} \times F_{7.9 \text{ dust}}$ , which defines the upper envelope of HII nuclei. Galaxies with no H<sub>2</sub> excess according to this definition were arbitrarily placed at an ordinate of 2.25 below the dashed line. The size of the symbols is proportional to the fraction of excess H<sub>2</sub> to total H<sub>2</sub> emission, coded as in Fig. 12. Sy nuclei are marked by an overlying horizontal bar, and star-forming nuclei by a square. The large scatter and the high ratios of excess H<sub>2</sub> to X-ray luminosities argue against X-rays playing a dominant role in the additional (nonstellar) excitation of H<sub>2</sub> in LINER/Sy nuclei (see text).

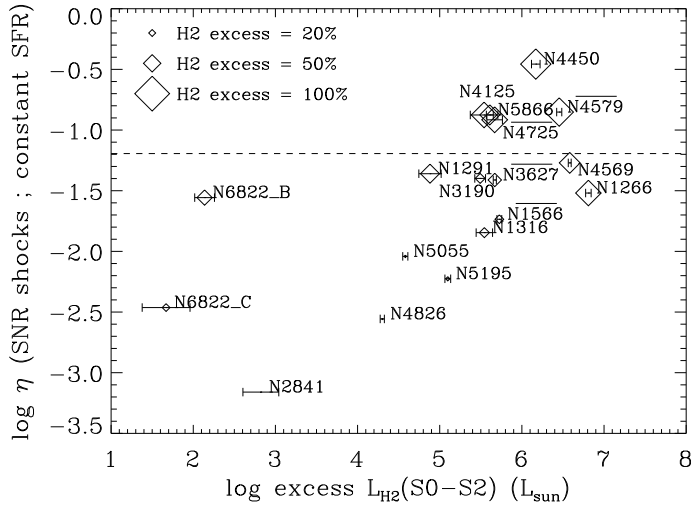


Fig. 12.— Heating efficiency required to account for the excess  $\text{H}_2$  emission by supernova remnant shocks (ratio of the power in the sum of the S(0)–S(2) lines to the total mechanical power), as a function of the excess  $\text{H}_2$  luminosity, defined as in Fig. 11. The dashed line indicates the maximum efficiency expected if all the mechanical power produced by supernova remnants is absorbed in molecular clouds, for the model parameters of Kaufman & Neufeld (1996). The size of the symbols is proportional to the fraction of excess  $\text{H}_2$  to total  $\text{H}_2$  emission. Sy nuclei are marked by an overlying horizontal bar. The regions B and C within NGC 6822 are the only non-AGN targets in this figure.