

White dwarf-open cluster associations based on *Gaia* DR2

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ABSTRACT

Context. Fundamental parameters and physical processes leading to the formation of white dwarfs (WDs) may be constrained and refined by discovering WDs in open clusters (OCs). Cluster membership can be utilized to establish the precise distances, luminosities, ages, and progenitor masses of such WDs.

Aims. We compile a list of probable WDs that are OC members in order to facilitate WD studies that are impractical or difficult to conduct for Galactic field WDs.

Methods. We use recent catalogs of WDs and OCs that are based on the second data release of the *Gaia* satellite mission (GDR2) to identify WDs that are OC members. This crossmatch is facilitated by the astrometric and photometric data contained in GDR2 and the derived catalogs. Assuming that most of the WD members are of the DA type, we estimate the WD masses, cooling ages, and progenitor masses.

Results. We have detected several new likely WD members and reassessed the membership of the literature WDs that had been previously associated with the studied OCs. Several of the recovered WDs fall into the recently reported discontinuity in the initial-final mass relation (IFMR) around $M_i \sim 2.0 M_\odot$, which allows for tighter constraints on the IFMR in this regime.

Key words. open clusters and associations; general – white dwarfs – catalogs – surveys

1. Introduction

White dwarfs (WDs) are the evolutionary endpoint of low- and intermediate-mass stars, which constitute a vast majority of all stars in the Galaxy. Their nature as compact and dense stellar remnants has been an important test bed for many areas of fundamental physics and stellar evolution theories. However, the study of WDs has been hampered by their low brightness, meaning that only observations of the closest objects could yield reliable results (see, e.g., Liebert 1980; Althaus et al. 2010; Córscico et al. 2019, for a general review).

White dwarfs associated with star clusters are extremely valuable. Star clusters are groups of gravitationally bound stars born in the same star-forming event, thus sharing the same age, metallicity, distance from the Sun, and proper motion. Since the WD cluster members also share these characteristics, this allows for a number of interesting questions to be addressed. Perhaps the most fundamental is the initial-final mass relation (IFMR), which links the final mass of a WD to the initial mass of its progenitor, hence also providing the total amount of mass lost during the stellar evolution. The progenitor mass can be estimated by determining the cooling age of a WD and subtracting it from the total age of the cluster as determined from the observations of the non-WD cluster members. This yields the lifetime of the WD progenitor, which can then be converted into the progenitor initial mass. Knowledge of the IFMR has applications in many areas of astrophysics. Perhaps one of the most fundamental applications of the high-mass end of the IFMR is determining the minimum main sequence stellar mass for a core-collapse supernova (SN) to occur. The IFMR is also an important ingredient in the modeling of stellar feedback in galaxy simulations and predicting SN type Ia rates (e.g., Greggio 2010; Agertz & Kravtsov

2015; Cummings 2017). Aside from the IFMR, other possible avenues of research utilizing cluster WDs include studying the effects of metallicity and binarity on WD evolution and measuring WD masses using gravitational redshift (Pasquini et al. 2019). Such studies are impossible or very challenging to conduct for Galactic field WDs.

While isolated WDs in globular clusters are very faint due to the considerable distances of these objects, the impetus for discovering WDs in open clusters (OCs) in the solar neighborhood is clear, as these OCs usually have well-determined parameters such as distance, reddening, age, and metallicity, providing a unique laboratory for studying the WDs associated with them and the related physical processes. This potential was realized early on when the Hyades cluster was studied by Tinsley (1974) and van den Heuvel (1975). More WD-cluster pairs were investigated by Weidemann (1977) and Romanishin & Angel (1980). Follow-up studies by Koester & Reimers (1981, 1985, 1993), Reimers & Koester (1982, 1988, 1989, 1994) obtained the spectroscopy of the WD candidates from Romanishin & Angel (1980), confirming some of them as bona fide cluster WDs and deriving their physical parameters. Since then, several other WD-OC pairs have been discovered and investigated by various authors and working groups (e.g., Anthony-Twarog 1982; Richer et al. 1998; Claver et al. 2001; Williams 2002). A recent compilation of OC WDs can be found in Cummings et al. (2018).

Past studies were limited by the small fields of view of the photometric surveys, which usually only covered the core OC regions. Another caveat was significant field WD contamination. To differentiate between the cluster and field WDs in the same area of the sky, accurate parallax and proper motion measurements of WDs were needed. The situation has

improved since the publication of the second data release of the *Gaia* mission (GDR2; [Gaia Collaboration 2016, 2018b](#)), which contains precise astrometry (positions, parallaxes, and proper motions) as well as photometry in three bands (G , G_{BP} , and G_{RP}). Since the advent of *Gaia*, the knowledge and census of Galactic OCs have also been substantially furthered (e.g., [Gaia Collaboration 2017](#); [Cantat-Gaudin et al. 2018a,b](#)). Furthermore, a large number of new WDs have been discovered and characterized ([Gentile Fusillo et al. 2019](#)), including WDs in OCs (e.g., [Salaris & Bedin 2018, 2019](#); [Richer et al. 2019](#)).

Due to recent increases in the number of known WDs and OCs with reliable parameters and astrometry, it has become possible to conduct a systematic search for WDs that are members of nearby OCs. In this paper, we crossmatch the known WDs and WD candidates listed in the catalog of [Gentile Fusillo et al. \(2019\)](#) with the OCs from [Cantat-Gaudin et al. \(2018a\)](#), using positional, parallax, and proper motion criteria. The physical reality of the putative WD-OC pairs are then further investigated using the cluster parameters (distance modulus, age, and reddening) and the position of the WD on the corresponding cooling sequence.

This paper is structured as follows. In Sect. 2, we describe the catalogs used in this study, the star cluster parameters, and the workflow leading to the selection of the WD OC member candidates. In Sect. 3, we discuss the recovered OCs hosting WDs and compare our detections with the literature, where available. The quality of GDR2 astrometric solutions and photometry for the recovered WDs are examined in Sect. 4. The WD masses and cooling ages are estimated in Sect. 5, and their application for the IFMR is addressed in Sect. 6. Finally, we summarize and add concluding remarks in Sect. 7.

2. Data analysis

The WD and OC catalogs that form the basis of this work are based on GDR2; therefore, they should be directly comparable, with no systematic offsets between them. The catalog of WD and WD candidates of [Gentile Fusillo et al. \(2019\)](#) lists over 480 000 objects, approximately 260 000 of which are high-probability WDs. Due to the intrinsic faintness of many isolated WDs, the majority of them are found within 1 kpc of the Sun, as can be seen in Fig. 1. This is in contrast with the distance distribution of the OCs from [Cantat-Gaudin et al. \(2018a\)](#); containing 1229 OCs), which is approximately uniform in the interval from 0.5 to 4 kpc; however, there is a notable paucity of OCs with distances ≤ 0.5 kpc. More than half of the cataloged WDs lie within this distance, with their distance distribution peaking at ~ 170 pc.

As can be seen in Fig. 1, the distribution of the parallax and proper motion errors of WDs and OC member stars listed in [Cantat-Gaudin et al. \(2018a\)](#) is also markedly different. The reason for this is two-fold. Firstly, the stars utilized to compute overall cluster astrometric parameters, which are also listed in [Cantat-Gaudin et al. \(2018a\)](#), are all brighter than $G > 18$ mag, whereas WDs from [Gentile Fusillo et al. \(2019\)](#) are much fainter by comparison, with a median brightness of $G_{WD} \approx 20$ mag. Such a jump in G leads to considerably larger errors for WDs ([Lindegren et al. 2018](#)). The second reason is that WDs are typically bluer in color than most stars in the GDR2. Blue objects observed by *Gaia* also exhibit increased errors in proper motion and parallax¹.

¹ <https://www.cosmos.esa.int/web/gaia/science-performance>

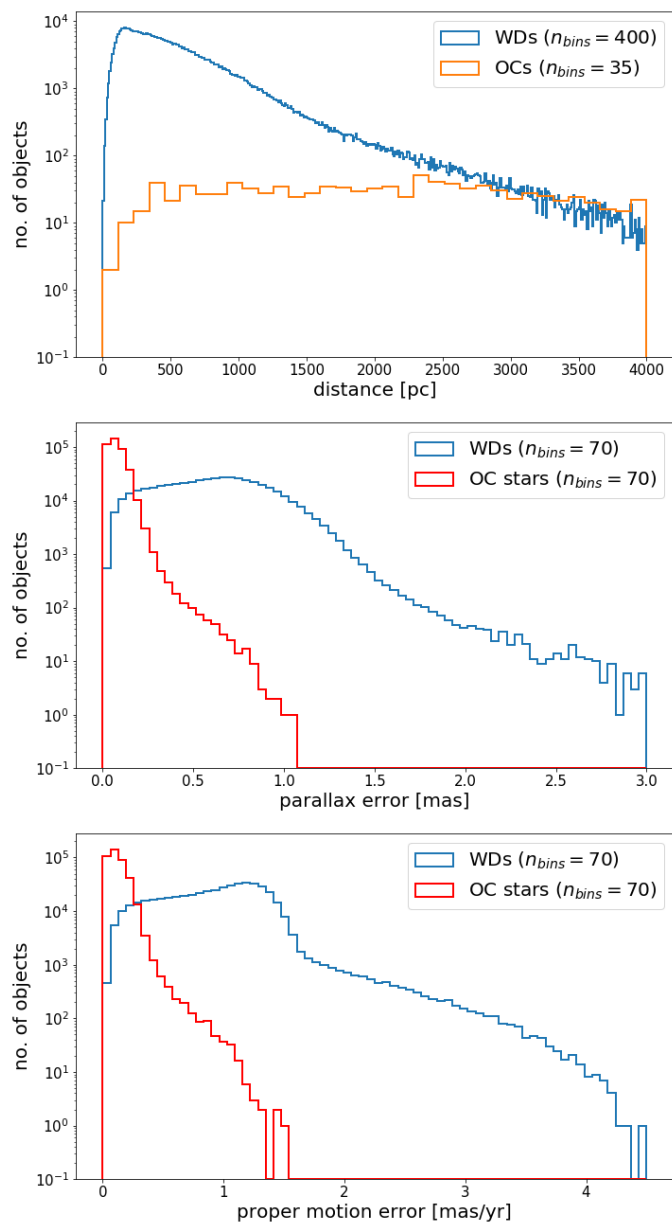


Fig. 1. *Top:* distance distribution of WDs from the catalog of [Gentile Fusillo et al. \(2019\)](#) compared to the distribution of OCs listed in [Cantat-Gaudin et al. \(2018a\)](#). *Middle:* comparison of the parallax error distribution of the WDs and OC member stars listed in [Cantat-Gaudin et al. \(2018a\)](#). *Bottom:* comparison of the average proper motion error (average of the RA and Dec components) of the WDs and OC members.

Due to these factors, using only the astrometric criteria (relying on positions, parallaxes, and proper motions) will yield a lot of low-confidence or spurious WD-OC matches. The most common such case is erroneous matches where a nearby WD gets matched with a more distant OC.

2.1. WD-OC pair preselection

Despite the shortcomings discussed above, the astrometric data are still potent when assigning potential WD members to OCs, especially when no such data of this quality and scope were available before GDR2. In order to make a rough preliminary preselection of potential cluster WDs, we utilized the

positional, proper motion, and parallax information contained in Cantat-Gaudin et al. (2018a) and Gentile Fusillo et al. (2019). The matching criteria are as follows:

$$\theta < 4.5 \times r50 \quad (1)$$

$$(plx - 3 \times s_plx; plx + 3 \times s_plx)_{OC} \cap (Plx - 3 \times e_Plx; Plx + 3 \times e_Plx)_{WD} \neq \emptyset \quad (2)$$

$$(pmRA - 3 \times s_pmRA; pmRA + 3 \times s_pmRA)_{OC} \cap (pmRA - 3 \times e_pmRA; pmRA + 3 \times e_pmRA)_{WD} \neq \emptyset \quad (3)$$

$$(pmDE - 3 \times s_pmDE; pmDE + 3 \times s_pmDE)_{OC} \cap (pmDE - 3 \times e_pmDE; pmDE + 3 \times e_pmDE)_{WD} \neq \emptyset. \quad (4)$$

Equation (1), where θ is the angular distance from a WD to a center of the cluster, represents the positional condition. Cantat-Gaudin et al. (2018a) list $r50$, which is the cluster radius that contains half of the cluster members, as the dimension of the studied clusters. In order to ensure search completeness, we considered WDs with projected separations up to $4.5 \times r50$ from the given cluster center. Next, Eq. (2) represents the parallax (distance) constraint. We considered every WD-cluster pair that satisfies this condition, where the WD has a parallax value of Plx and an associated error e_Plx from Gentile Fusillo et al. (2019; adopted directly from the GDR2) and the OC has a mean parallax of plx and a standard deviation of parallax of OC members s_plx from Cantat-Gaudin et al. (2018a). Lastly, Eqs. (3) and (4) are proper motion constraints. Again, Gentile Fusillo et al. (2019) adopt proper motion values and errors directly from the GDR2. For OCs, $pmRA$ ($pmDE$) is the mean proper motion along the right ascension (declination) of OC members, and s_pmRA (s_pmDE) is its standard deviation.

Such a selection yields almost 4000 distinct WD-OC pairs. Naturally, due to the problems with the WD astrometry outlined in Sect. 2 and the generous selection criteria applied, most of these pairs are low-probability and are only spurious pairings. Given the nature of the WD astrometry, it is normally not sufficient to rely on astrometric data alone to determine membership. Further investigations can be conducted using cooling models in conjunction with cluster ages.

2.2. Isochrones and white dwarfs

One of the most important parameters describing stellar clusters is their age. With the use of photometric data available for the cluster members, the age of the cluster is usually found with the help of an isochrone fitting method. First, isochrones need to be calculated, which can be done with evolutionary models for stars of different masses. In the case that a correct age and metallicity are chosen (together with the distance and the extinction), the resulting isochrone should coincide with the distribution of cluster members in the color-magnitude diagram (CMD). Due to its dependence on all four cluster parameters, this method is very useful for improving distance and extinction while determining age and metallicity (although metallicity is often ignored and assumed to be solar). This whole process is a necessary step because of the fact that we are attempting to assign WDs to clusters. In this section, our goal is to show the quality of the cluster parameters derived from isochrone fitting techniques that have (mostly) been published in recent years. Furthermore, the method used to compute values for the WDs displayed in the CMD (in *Gaia* magnitudes) is described.

Table 1. Sources for cluster parameters.

Source of parameters	Number of OCs
Bossini et al. (2019)	67
Kharchenko et al. (2013)	81
Dias et al. (2002)	2
Röser et al. (2016)	3
Custom fit	98

To verify our assignment of WDs to the sample of OCs, we need to take a look at the CMDs that show both the cluster members and the WDs. Moreover, we need to acquire cluster parameters (distance, extinction, and age, excluding metallicity) for all clusters in our sample. The newest data set provided by Bossini et al. (2019) contains the required parameters for 269 clusters, which are based on the data from the GDR2. Unfortunately, not all of these clusters coincide with those from our sample. For this reason, we decided to also make use of the data provided by Kharchenko et al. (2013) We took parameters from Dias et al. (2002) and Röser et al. (2016) as secondary sources of data if a cluster is not present in either of the two previous data sets.

Closer inspection of the individual CMDs then helped us determine which of the data sets gives a better isochrone fit to a given cluster. For our purposes, we decided to use CMD 3.3, the isochrone data from Evans et al. (2018), an assumed solar metallicity ($Z = 0.02$), and a chosen time-step $\Delta \log T = 0.05$. We favored this metallicity value because it has been shown to be consistent with recent results of helioseismology (Vagnozzi 2019). Together with information about cluster members from Cantat-Gaudin et al. (2018a) and the sets of cluster parameters, we can make a comparison between the corresponding isochrones. It is immediately clear from the plots that many of the clusters were assigned parameters that correspond to isochrones that do not match these clusters well enough. Our criterion for picking the parameters from the available data was to get the best isochrone fit. For the most part, values from Bossini et al. (2019) and Kharchenko et al. (2013) provide the best descriptions of the clusters (for example, Fig. 2), with parameters of only five clusters being taken from the secondary data sets. However, there are also many examples (about one-third of the whole sample) of clusters for which it was impossible to get an acceptable fit using data from any of the mentioned works.

For these cases, we fit the isochrones of all the individual clusters, using the photometric data of stars with membership probabilities larger than 50%. This was done without any black box algorithm. The metallicity was again assumed to be solar and kept fixed. Then, the reddening was determined using the shape of the main sequence. As a last step, the distance modulus was chosen so that the main sequence and turnoff point fit satisfyingly within the isochrone grid. The total final result for cluster parameters can be seen in Table 1.

The next task was fairly simple: determine the position of the WDs in the CMDs. To do this properly, we had to be able to subtract the extinction from the *Gaia* magnitudes. Since the extinction is usually described by either the A_V or R_V parameters (we assumed that $A_V = \frac{E(B-V)}{0.324}$) and we want to make use of GDR2 data, we needed to know the transformations between extinction in A_G (A_{BP} , A_{RP}) and A_V . It is not viable to use the simple approach $A_G = 0.835 A_V$ due to the width of the *Gaia*

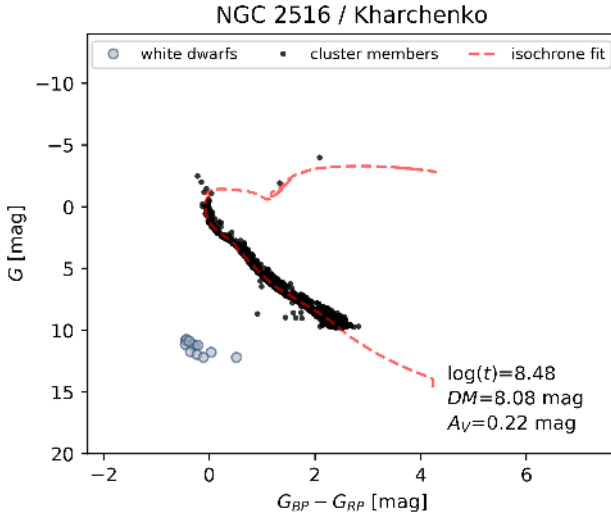


Fig. 2. Example of a cluster (NGC 2516) in the CMD with members taken from [Gentile Fusillo et al. \(2019\)](#) and fit with an isochrone (parameters from [Kharchenko et al. 2013](#)). Our initial candidate WDs are displayed in the plot together with the cluster parameters (age, reddened distance modulus, and extinction).

passbands. For our purposes, we decided to use the polynomial combination of $(G_{BP} - G_{RP})$ and A_V values that is described in [Gaia Collaboration \(2018a\)](#).

As mentioned before, we only employed isochrones with solar metallicity (i.e., $Z = 0.02$). To investigate the effect of the metallicity on the cluster parameters derived from isochrone fitting, the range of the metallicity in the solar vicinity has to be assessed. [Netopil et al. \(2016\)](#) present homogenized metallicities for 172 OCs on the basis of photometric and spectroscopic data. More recent studies using optical ([Pancino et al. 2017](#)) or infrared ([Donor et al. 2018](#)) spectroscopy have not added a significant number of new investigated OCs. Furthermore, these results are very much in line with those from [Netopil et al. \(2016\)](#). These last authors have showed that almost all OCs within 2 kpc of the Sun have $[\text{Fe}/\text{H}] = \pm 0.2$ dex. There are hardly any known Galactic OCs that exceed a $[\text{Fe}/\text{H}]$ value of ± 0.5 dex. The isochrones up to $[\text{M}/\text{H}] = \pm 1.0$ dex are shifted in the distance modulus only. This means that, for the same color, stars become fainter for lower metallicities. We used the turnoff points for the whole isochrone grid to investigate the concrete values. As a conclusion, it can be said that for $[\text{M}/\text{H}]$ up to ± 1.0 dex, the differences of the distance modulus scales are one-to-one with metallicity (i.e., $\Delta[\text{M}/\text{H}] \approx \Delta DM$). This shift is negligible compared to the width of the main sequence and the intrinsic errors of the parallaxes. Therefore, using an isochrone grid with solar metallicity is a justifiable approach.

Finally, we wanted to compare the calculated extinction values A_G with those provided by [Gentile Fusillo et al. \(2019\)](#). Assuming that $A'_G = 0.835 A_V$, they give

$$A_G = A'_G \left(1 - \exp\left(-\frac{\sin |b|}{200\varpi}\right) \right),$$

$$A_{(BP-RP)} = 0.586 A'_G \left(1 - \exp\left(-\frac{\sin |b|}{200\varpi}\right) \right),$$

as the effective values of the extinction coefficients, where b is the Galactic latitude of the WD and ϖ is its parallax (in arcseconds). We can see that the relation between the two results is not one-to-one (Fig. 3). However, this is to be expected since both

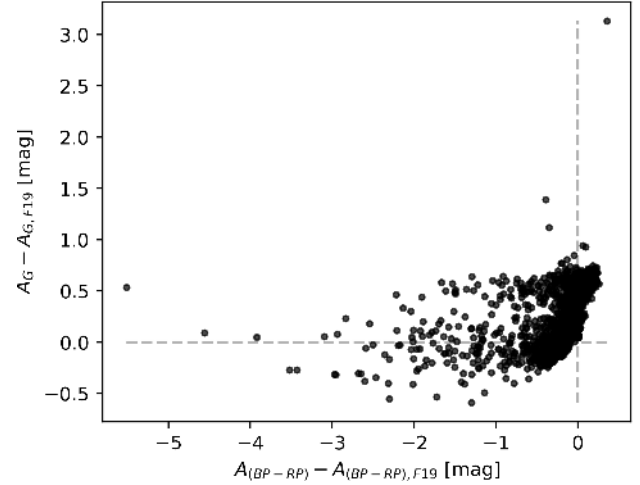


Fig. 3. Comparison of the extinction values A_G and A_{BP-RP} between this work and [Gentile Fusillo et al. \(2019\)](#).

approaches use a different version of the extinction law. What remains uncertain in our case is the applicability of the transformation described in [Gaia Collaboration \(2018a\)](#) since their coefficients were derived with the use of stars with estimated effective temperatures $T_{\text{eff}} \lesssim T_{\text{eff,WD}}$ and it is unknown what order of magnitude of errors is produced at the higher temperature regime ($> 10\,000$ K).

2.3. CMD and cooling age-based filtering

Provided that accurate cluster ages, distances (parallaxes), and extinction values are available, it is possible to use photometry to filter out spurious WD-OC pairings. In order to do this, we used the cluster parameters as obtained in the previous section and Montreal WD cooling tracks² ([Fontaine et al. 2001](#)).

For our initial sample of several hundred putative WD-OC pairings, we used the distance moduli and extinctions of the matched OCs to compute the dereddened absolute magnitudes and colors for the corresponding WDs. We plot these quantities with the theoretical cooling tracks for the lowest- ($0.2 M_{\odot}$) and highest-mass ($1.2 M_{\odot}$) WDs in Fig. 4. In order for a WD-OC pairing to be physical (provided that the WD is not in a binary), it is necessary (but not sufficient) for a WD to lie in the CMD region delineated by the lowest- and highest-mass cooling tracks. It is apparent that the majority of the potential OC WDs lie above the lowest-mass cooling track, being more luminous than what would be expected if they were OC members. This was expected (see the discussion in Sect. 2), as these WDs tend to be in the foreground of the OCs and are spuriously matched to them due to the generous selection criteria and substantial errors in parallax and proper motions.

Further constraints can be made using the age of the OC matched with a WD. Obviously, the cooling age of the WD cannot be higher than the age of the OC it is associated with, provided that the association is real. Using this, other spurious WD-OC pairs can be filtered out on an individual basis using additional cuts in the CMD diagrams. If the cluster age is known, a WD that is associated with the cluster should lie in the CMD region delineated by the lowest- and highest-mass cooling track (as discussed above), the zero-age cooling isochrone, and the cooling isochrone corresponding to the cluster age.

² <http://www.astro.umontreal.ca/~bergeron/CoolingModels>

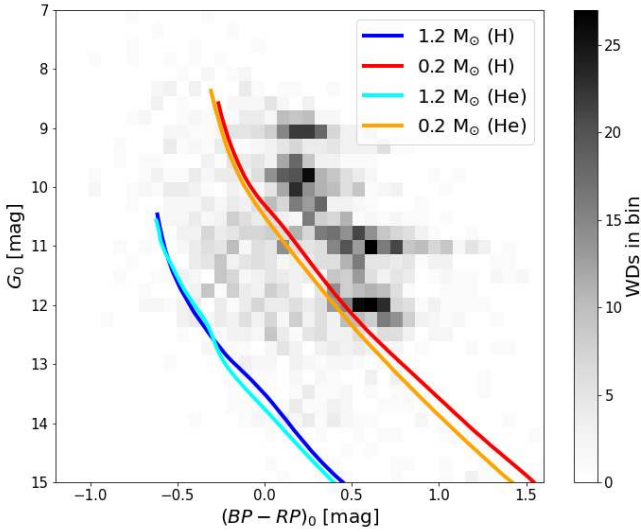


Fig. 4. 2D density WD histogram obtained from the initial WD sample in the absolute magnitude-color space. The absolute magnitudes and colors for each WD are calculated using the parameters of the cluster of which the WD is a member candidate. Overlaid are the Montreal WD cooling tracks for low-mass and high-mass WDs with H and He atmospheres.

3. Notes on the individual WD-OC pairs

In this section, we list and discuss the obtained OC-WD candidate pairs that passed the astrometric, photometric, and cooling age criteria as described in the previous sections. The figures that illustrate the placement of the WD candidates in the cluster CMD and astrometric phase space are included in the appendix; however, some of the more interesting examples are discussed in this section.

3.1. ASCC 73, ASCC 79, and ASCC 97

ASCC 73, ASCC 79, and ASCC 97 are three OCs discovered in [Kharchenko et al. \(2005\)](#). Due to their relatively recent discovery and sparse nature, they have been studied very little in the literature. No studies of WDs potentially hosted by these clusters have been conducted to date.

Our analysis has recovered one potential cluster WD: GDR2 5856401252012633344 for ASCC 73. On face value, it seems to be a mild outlier from the other cluster members as cataloged by [Cantat-Gaudin et al. \(2018a\)](#), both in terms of proper motion and parallax. However, considering the astrometric uncertainties of the WD candidate, it is still consistent with cluster membership.

For ASCC 79, we have found three possible cluster WDs: GDR2 5825203021908148480, 5826384584601681152, and 5825187834899772160. However, it needs to be noted that the probability of the last object being a WD, as given in [Gentile Fusillo et al. \(2019\)](#), is only $P_{wd} = 0.59$.

Gaia DR2 4092407537313874048 has been identified as a viable candidate for ASCC 97. While its astrometric properties are consistent with cluster membership, its WD nature is ambiguous ($P_{wd} = 0.47$ in [Gentile Fusillo et al. 2019](#)).

3.2. Alessi 3

Alessi 3 is a sparse evolved OC (or OC remnant; [Angelo et al. 2019](#)). Its WD content has not been studied before.

We have identified one cluster WD candidate: GDR2 5508976051738818176. Its astrometric properties are consistent with cluster membership, but its parallax puts it into the cluster background if taken at face value. However, the parallax error is very high, and a number of cluster members lie within 1σ of the cluster WD candidate's parallax.

3.3. Alessi 13

Alessi 13 (χ^{01} For moving group) is a sparse nearby stellar association. Its WD content has never been studied.

We have identified one possible WD cluster member: GDR2 4853382867764646912. Its astrometric properties are consistent with cluster membership.

3.4. Alessi 62

Alessi 62 is another unstudied old OC. No WDs that are potential members of this cluster are known.

Our analysis has yielded one cluster WD candidate: GDR2 4519349757791348480. Its proper motion is consistent with cluster membership; however, its parallax is more problematic as it suffers from a large uncertainty, and, if taken at face value, it puts the member candidate into the background. However, some of the cluster members are still contained within its 1σ uncertainty interval. Its nature as a bona fide WD is ambiguous since [Gentile Fusillo et al. \(2019\)](#) gives a lower $P_{wd} = 0.56$ for this object.

3.5. IC 4756

IC 4756 is a close, intermediate-age OC. Though IC 4756 has been heavily studied, WDs potentially hosted in the cluster have never been investigated in detail in the literature. However, it needs to be noted that by looking at the CMD of the cluster stars listed in [Cantat-Gaudin et al. \(2018a\)](#), one can readily identify a potential WD candidate on the cluster WD sequence. The WD is bright enough to not be excluded in the magnitude cutoff of $G = 18$ mag adopted there.

Our analysis has identified only one viable cluster WD candidate, and it is the same one as discussed above (GDR2 4283928577215973120). Its proper motion and parallax make it a very likely cluster member.

3.6. Mamajek 4

Mamajek 4 is a poorly studied OC. No WD studies targeting this cluster have been conducted.

Our search has identified one potential cluster WD: GDR2 6653447981289591808. Its proper motion is consistent with cluster membership, though its parallax indicates that it may be a background object. However, its parallax error is quite high and a significant portion of the cluster members lie within a 1σ error of the candidate parallax.

3.7. Melotte 22

Melotte 22 (Pleiades) is one of the closest, best-studied, and, arguably, most well-known OCs. Despite its proximity, only one cluster WD has been identified so far: EGGR 25 (GDR2 66697547870378368; [Eggen & Greenstein 1965](#); [Lodieu et al. 2019](#)).

Our analysis recovered EGGR 25. However, it failed to identify any new potential cluster WD candidates.

3.8. NGC 2422

NGC 2422 is a rather young (~ 150 Myr) OC with a current turnoff age of about $5.4 M_{\odot}$ (Richer et al. 2019). The potential WD content of the cluster was first investigated by Koester & Reimers (1981), who found a potential WD candidate (GDR2 3030026344167186304) that may also be a cluster member. However, they were not able to fully ascertain its nature; while it may be a massive WD that is a member of the cluster, it may also be a field WD behind the cluster or a subdwarf O-type star. Richer et al. (2019) find a massive cluster WD with a helium-rich atmosphere and large magnetic field, probably in a binary with a late-type companion (GDR2 3029912407273360512).

Our analysis only recovered the WD found by Richer et al. (2019), as the other one is not included in the catalog by Gentile Fusillo et al. (2019). However, taking advantage of GDR2 astrometry, it can clearly be seen that the WD member candidate of Koester & Reimers (1981) is most certainly not a cluster member and that it lies in the foreground.

3.9. NGC 2516

NGC 2516 is also a young OC that likely started forming WDs relatively recently. Reimers & Koester (1982) first identified three probable cluster WDs and later added a fourth, the nature of which was previously uncertain (Koester & Reimers 1996). Recently, Holt et al. (2019) have added two more candidate WD members, which were identified using the GDR2.

Our analysis of this cluster identified three sources, one of which was already identified in Reimers & Koester (1982) and the two others in Holt et al. (2019). Thus, no novel detections were made. The other three WDs from Reimers & Koester (1982) and Koester & Reimers (1996) are also included in Gentile Fusillo et al. (2019), but their cluster membership is not solid. GDR2 5290720695823013376 seems to lie in the foreground and GDR2 5290719287073728128 in the background; GDR2 5290834387897642624 is a proper motion outlier but just narrowly did not make the cut.

3.10. NGC 2527

NGC 2527 is an older (~ 800 Myr) OC with a turnoff mass of $\approx 2.2\text{--}3.5 M_{\odot}$ (Raddi et al. 2016). A WD that is also a likely cluster member was reported in Raddi et al. (2016).

We did not recover this WD (GDR2 5597874285564810880) as it is not listed in Gentile Fusillo et al. (2019). However, we identified a new candidate. Using the GDR2 astrometry, it can clearly be seen that the WD identified as a cluster member in Raddi et al. (2016) is a significant outlier in both parallax and proper motion, making it a likely field object.

3.11. NGC 2632

NGC 2632 (Praesepe) is a close and well-known OC with a large number of published WDs. It is considered to be a “benchmark” cluster for WD studies, and it is likely that the observed cluster single WD population is complete due to its proximity.

Our analysis recovered all 12 known cluster WDs with no new detections, as expected. A comprehensive analysis of these WDs in the context of their parent cluster is available in a recent analysis by Salaris & Bedin (2019) and the references therein.

3.12. NGC 3532

This rich, ~ 300 Myr old OC is believed to host a number of WDs. Reimers & Koester (1989) identified seven candidate cluster WDs and confirmed the degenerate nature of three of them. Their subsequent extended survey added three more candidate WD members later on (Koester & Reimers 1993). However, a more detailed analysis by Dobbie et al. (2009) put two of these WDs in the background of the cluster. An expanded survey by Dobbie et al. (2012) identifies several more WD candidates, including another four bona fide WDs in the direction of the cluster, three of which are reportedly cluster members. Furthermore, Raddi et al. (2016) add an additional, very massive WD cluster member.

A combined tally of seven cluster WDs, as obtained from the literature, makes the cluster appealing as one of the benchmark clusters, together with Hyades and Praesepe. However, our detection of only three WD candidate members is seemingly at odds with these reported WD numbers. Crossmatching these literature WDs with the GDR2 and querying them in the WD catalog by Gentile Fusillo et al. (2019), we found that only two of them are listed there: GDR2 5340219811654824448 and GDR2 5338718261060841472; the latter is also a cluster member according to our analysis. Our second identified cluster WD candidate is also among the cluster members reported in Cantat-Gaudin et al. (2018a) – GDR2 5340220262646771712 – with a reported membership probability of 1.0; it actually lies at the beginning of the WD cooling sequence. This makes it a solid WD candidate that must have formed very recently. The last detected source – GDR2 5338685962923467136 – is a new candidate cluster WD.

All of the reported cluster WDs, with the exception of the massive WD identified in Raddi et al. (2016), have a GDR2 counterpart with a full five-parameter solution. Despite them not being in the catalog of Gentile Fusillo et al. (2019), we can still assess their cluster membership. Figure 5 shows that the literature WDs have astrometric properties that are consistent with the cluster membership. The only exception is GDR2 5340148691289324416 (reported as a member in Dobbie et al. 2012), whose cluster membership, which is based on its astrometric properties, can be disputed. Another interesting case is GDR2 5338650984675000448 (cluster member according to Reimers & Koester 1989; also listed in Gentile Fusillo et al. 2019), which seems too luminous and red to be a cluster member.

3.13. NGC 6633

NGC 6633 is a loose OC with various age estimates, ranging from 430 Myr (Dias et al. 2002) to 800 Myr (our estimate from isochrone fitting). Reimers & Koester (1994) investigated possible WD candidates in the field of the cluster and found one (GDR2 4477214475044842368) that may be a cluster member, but they were not able to confirm its cluster membership. A later study by Williams & Bolte (2007) found two more WDs at the cluster distance modulus (GDR2 4477166581862672256 and GDR2 4477253202776118016) and another two (GDR2 4477214475044842368 and GDR2 4477168746525464064) that appear too bright to be cluster members if single, but could potentially be double degenerate systems belonging to the cluster. One of them had already been identified as a WD member candidate in Reimers & Koester (1994).

Our analysis yielded two WD member candidates: GDR2 4477214475044842368 and GDR2 4476643725433841920; one

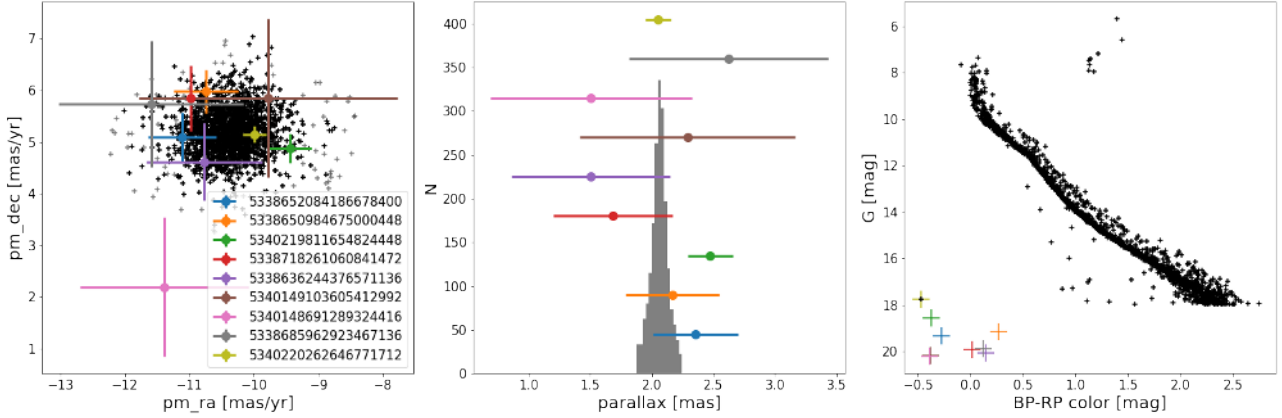


Fig. 5. *Left:* proper motion diagram of the NGC 3532 stars with the recovered and literature WD proper motion overlaid. Cluster stars with cluster membership probability <0.5 are marked using gray crosses, while black crosses indicate likely cluster members. Here, and in the subsequent graphs, the errorbars indicate a 1σ uncertainty, as reported in the GDR2. Errorbars for the cluster stars are omitted for clarity. *Middle:* parallax histogram of the cluster member stars (membership probability ≥ 0.5) with WD parallaxes overlaid. *Right:* cluster member star CMD with WDs overlaid.

was already known and one is a novel detection. Out of the two WD member candidates identified in Williams & Bolte (2007), we identified one as a cluster member in our analysis. Neither of them is included in Gentile Fusillo et al. (2019). *Gaia* DR2 4477166581862672256 has a parallax and proper motion consistent with cluster membership. The other, which was thought to be a rare DB cluster WD, is a clear outlier in terms of both parallax and proper motion. Out of the two potential double degenerate systems (both listed in Gentile Fusillo et al. 2019), only one of them (GDR2 4477214475044842368) has astrometric parameters consistent with cluster membership.

3.14. NGC 6991

NGC 6991 is a relatively unstudied sparse OC. Our literature search for cluster WDs and candidates did not yield any objects that may be associated with this cluster.

We present the identification of a possible cluster WD (GDR2 2166915179559503232). It is a high-confidence WD in Gentile Fusillo et al. (2019), and its proper motion is consistent with other members of the cluster. On face value, its parallax puts it in the foreground of NGC 6991, but the parallax error is rather large so its cluster membership cannot be conclusively assessed this way.

3.15. NGC 7092

NGC 7092 (M 39) is a well-known and well-studied cluster. At the time of writing, Caiazzo et al. (2020) have identified and characterized one cluster WD (GDR2 2170776080281869056).

Our analysis yielded a high-confidence WD that is a possible member of this cluster, the same object as in Caiazzo et al. (2020). The parallax and proper motion of this object matches well with those of the cluster members.

3.16. RSG 7 and RSG 8

RSG 7 and RSG 8 are two of the sparse, close OCs discovered in Röser et al. (2016). The literature on these clusters is very limited, and there are no WDs associated with them.

Our search resulted in three WD candidates that can potentially be assigned to RSG 7, as well as one that could be a

member of either RSG 7 or RSG 8 (the double match resulted from a combination of the close proximity of the clusters in the projection on the sky as well as the proper motion space and large parallax uncertainty of the member candidates). However, upon analysis of the proper motion diagram, parallax distribution, and CMDs of the cluster members in Cantat-Gaudin et al. (2018a), we concluded that the parameters of these clusters listed there are erroneous. The issue seems to be a heavy contamination from the members of the adjacent cluster, which is clearly visible and presents as multiple populations in the cluster CMDs. Taking the quality of the astrometric parameters of the candidate WD members into consideration as well, we thus discarded these OC-WD pairs.

3.17. Ruprecht 147

Ruprecht 147 (NGC 6774) is one of the oldest star clusters in the solar neighborhood. Its proximity and age make it attractive as one of the potential benchmark clusters for stellar evolution studies, and WDs in particular. This has been demonstrated by Gaia Collaboration (2018a), who identified ten cluster WDs. A subsequent comprehensive study by Olivares et al. (2019) has added five more, for a total WD tally of 15. A recent study of the cluster by Marigo et al. (2020) rules out the membership of several previously associated WDs based on conflicting spectroscopic and photometric luminosities, but it adds one new cluster WD not listed in Gentile Fusillo et al. (2019)

Our analysis identified nine cluster WD candidates, none of which are new detections; this is not surprising given the depth of the previous studies. However, we decided to discard three member candidates – GDR2 4183847562828165248, GDR2 4184148073089506304, and GDR2 4184196073644880000 – which are all members according to Gaia Collaboration (2018a) and Olivares et al. (2019) but are found to be non-members in Marigo et al. (2020). Therefore, we retained six potential WD members. One of the WDs from Olivares et al. (2019) is not included in Gentile Fusillo et al. (2019), and five of them are slight proper motion outliers with respect to the cluster members of Cantat-Gaudin et al. (2018a), with one of the WDs just narrowly inside the margin delineated by our selection criteria.

3.18. Stock 2

Stock 2 is a nearby OC. Despite its proximity, it is relatively unstudied due to its large angular size and the variable reddening in its direction (Spagna et al. 2009). Its age is disputed, so we estimated the cluster age to be $\log(t) = 8.5$. Stock 2 was one of the clusters studied in Gaia Collaboration (2018a), who identify eight cluster WD candidates.

Our analysis managed to identify 16 WD candidates with parameters consistent with cluster membership. Out of these, ten are new detections, while the remaining six were identified in Gaia Collaboration (2018a). There are two extra cluster WD candidates contained in Gaia Collaboration (2018a) that were not recovered in our analysis, despite them being listed in Gentile Fusillo et al. (2019): GDR2 508400329710144896 and GDR2 506848643933335296. The parallaxes of these two objects are not consistent with cluster membership.

3.19. Stock 12

Stock 12 is a poorly studied cluster, the WD content of which has never been studied before. We uncovered only one novel WD member candidate: GDR2 1992469104239732096.

4. Reliability of the GDR2 solution

The GDR2 provides high-quality astrometric and photometric measurements for an unprecedented number of sources. However, it still contains some solutions that are ill-behaved and need to be accounted for or removed from the analysis. Problems with the astrometry and photometry can arise for sources that are located in regions with high source densities, for instance in the Galactic plane and star clusters. Binary systems can also be problematic because GDR2 sources are treated as single stars in the astrometric solution, whereas binaries do not receive any special treatment (Gaia Collaboration 2018b; Lindegren et al. 2018). We therefore examined the quality of the GDR2 solutions for the recovered WD member candidates.

Gentile Fusillo et al. (2019) have conducted some cleaning of their WD sample, identifying many potentially spurious sources. However, in order to obtain a reliable list of WD member candidates, we further cleaned the WD sample based on the recommended astrometric and photometric flags. Informed by Gaia Collaboration (2018b), Lindegren et al. (2018) and Lindegren (2018; GAIA-C3-TN-LU-LL-124-01³), we retained the sources that satisfied the following three conditions: (a) duplicated_source = False; (b) astrometric_excess_noise < 1 mas or astrometric_excess_noise_sig < 2; and (c) ruwe < 1.4.

Specifically, the flag duplicated_source=True implies observational problems, crossmatching problems, processing problems, or stellar multiplicity, potentially leading to problems in the astrometric solution. The astrometric_excess_noise (ϵ_i) is the excess astrometric noise of the source postulated to explain the scatter of residuals in the astrometric solution. When it is high and significant, it can mean that the astrometric solution has failed for that source. Another possibility is that the observed source is a binary system, where the additional scatter can arise from the movement of the emission centroid due to the motion of the binary components. Finally, the cuts based on ruwe, which stands for renormalized unit weight error, ensured the removal of ill-behaved astrometric solutions.

None of the selected WD candidates exhibited increased astrometric noise or ruwe values; however, three of them (GDR2 4519349757791348480, GDR2 5338685962923467136, and GDR2 511159317926025600) were possible duplicated sources. These objects were then removed from the candidate list.

In order to identify the cases where the photometry is unreliable, we applied the following two quality indicators, as given in Gaia Collaboration (2018a): (a) $\text{phot_bp_rp_excess_factor} > 1.0 + 0.015(G_{BP} - G_{RP})^2$ and (b) $\text{phot_bp_rp_excess_factor} < 1.3 + 0.06(G_{BP} - G_{RP})^2$. The WDs that did not satisfy the above criteria were retained as member candidates, but we did not estimate their characteristics as the photometry cannot be considered reliable.

5. Parameter estimates for the recovered WD member candidates

In order to establish precise WD parameters, spectroscopic studies are usually needed. In addition to atmospheric parameters such as effective temperature, surface gravity, and chemical composition, spectroscopic data provide an additional check for cluster membership by comparing the WD spectroscopic-based luminosity with the luminosity derived from photometry when the cluster distance and extinction is adopted. Furthermore, spectroscopy is required to ascertain the WD atmospheric composition (unless ultraviolet photometry is available) and binarity status. Unfortunately, most of the new WD member candidates lack the needed spectroscopic data. However, we can assume that most of the recovered WDs are of the DA type, which is overwhelmingly the most dominant WD type found in OCs due to their typical ages, while only a handful of DB cluster WDs are known in the literature (e.g., Kalirai et al. 2005; Salaris & Bedin 2019; Marigo et al. 2020). Under this assumption, the GDR2 photometry enables us to compute the WD absolute magnitudes and colors, adopting the cluster distance and reddening. From these, the photometric-based estimates of WD parameters, such as mass M_{WD} and cooling age t_{cool} , can be derived.

While the Montreal WD cooling tracks were used for the photometric selection of viable OC WDs and can, in principle, be used to compute M_{WD} and t_{cool} estimates, they suffer from several shortcomings that can affect these estimates. Notably, they do not include the effects of residual nuclear burning, which can have a significant impact on the derived t_{cool} (Iben & Tutukov 1984; Camisassa et al. 2015; Althaus et al. 2010). Additionally, the Montreal WD cooling tracks assume unrealistic WD core compositions and do not include the impact of the energy release resulting from phase separation on crystallization, which also affects the derived t_{cool} . Then, to compute M_{WD} and t_{cool} , we used a combination of models, employing the tool from Cheng (2020). For the WDs with masses of $0.45 M_{\odot} \lesssim M_{WD} \lesssim 1.0 M_{\odot}$, we used the model from Renedo et al. (2010) with a metallicity of $Z = 0.01$, which is suitable for the solar neighborhood. For the high-mass WDs ($M_{WD} \gtrsim 1.0 M_{\odot}$), we adopted the model from Camisassa et al. (2019), in which such WDs are expected to be harboring O-Ne cores. In order to account for the errors in absolute magnitude and color, we performed a 10^4 -element Monte Carlo simulation for each WD, interpolating the M_{WD} and t_{cool} from the cooling tracks each time. For the simulations, we drew absolute magnitude and color samples from normal distributions (assumed to be independent), which are centered around the measured values and 1σ errors. We defined our 1σ absolute magnitude and color errors by adding in quadrature the error from the distance modulus (in the case of absolute magnitude), reddening, and instrumental errors. Resulting

³ http://www.rssd.esa.int/doc_fetch.php?id=3757412

Table 2. Novel or newly characterized WD-OC pairs recovered in this analysis.

GDR2 source ID	Associated cluster	P_{WD}	$\log t_{\text{cl}}$ (yr)	[Fe/H]	[M/H]	M_{WD} (M_{\odot})	t_{cool} (Gyr)
5856401252012633344	ASCC 73	0.867	8.190			$0.64^{+0.12}_{-0.11}$	$0.097^{+0.037}_{-0.03}$
5825203021908148480	ASCC 79	0.961	6.950			$0.37^{+0.13}_{-0.08}$	$0.007^{+0.004}_{-0.004}$
5826384584601681152	ASCC 79	0.916	6.950			$0.33^{+0.09}_{-0.06}$	$0.008^{+0.002}_{-0.005}$
5825187834899772160	ASCC 79	0.594	6.950			$0.29^{+0.07}_{-0.03}$	$0.01^{+0.002}_{-0.003}$
4092407537313874048	ASCC 97	0.465	7.900		0.129 ± 0.166	$0.24^{+0.06}_{-0.04}$	$0.035^{+0.023}_{-0.01}$
5508976051738818176	Alessi 3	0.995	8.870		-0.275 ± 0.065	$0.81^{+0.09}_{-0.09}$	$0.638^{+0.128}_{-0.109}$
4853382867764646912	Alessi 13	0.998	8.720		0.06 ± 0.15	$0.57^{+0.08}_{-0.08}$	$0.568^{+0.076}_{-0.07}$
4283928577215973120	IC 4756	0.986	8.987	-0.02 ± 0.01		$0.34^{+0.14}_{-0.07}$	$0.011^{+0.005}_{-0.006}$
6653447981289591808	Mamajek 4	0.990	8.824	0.09 ± 0.08		$0.85^{+0.11}_{-0.12}$	$0.282^{+0.073}_{-0.061}$
5289447182180342016	NGC 2516 ^(a)	0.999	8.475	0.08 ± 0.01		$0.71^{+0.21}_{-0.17}$	$0.149^{+0.069}_{-0.052}$
5294015515555860608	NGC 2516 ^(a)	0.998	8.475	0.08 ± 0.01		$0.98^{+0.11}_{-0.11}$	$0.077^{+0.027}_{-0.025}$
5597682038533250304	NGC 2527	0.996	8.910	-0.1 ± 0.04		–	–
5340220262646771712	NGC 3532	0.989	8.650	-0.07 ± 0.10		$0.5^{+0.12}_{-0.12}$	$0.3^{+0.061}_{-0.063}$
4476643725433841920	NGC 6633	0.532	8.900	-0.098 ± 0.037		$0.58^{+0.17}_{-0.16}$	$0.157^{+0.061}_{-0.055}$
2166915179559503232	NGC 6991	0.998	9.100	0.0 ± 0.03		$0.56^{+0.14}_{-0.12}$	$0.023^{+0.013}_{-0.012}$
418392888026931328	Ruprecht 147	0.996	9.330	0.16 ± 0.08		$0.49^{+0.27}_{-0.19}$	$0.162^{+0.091}_{-0.078}$
4183926006112672768	Ruprecht 147	0.955	9.330	0.16 ± 0.08		$0.49^{+0.11}_{-0.12}$	$0.481^{+0.076}_{-0.066}$
506514907785623040	Stock 2	0.939	8.500	-0.06 ± 0.03		$0.37^{+0.11}_{-0.08}$	$0.306^{+0.046}_{-0.045}$
508276703371724928	Stock 2	0.980	8.500	-0.06 ± 0.03		$0.39^{+0.48}_{-0.19}$	$0.169^{+0.159}_{-0.117}$
507054806657042944	Stock 2	0.999	8.500	-0.06 ± 0.03		$0.83^{+0.07}_{-0.08}$	$0.069^{+0.023}_{-0.018}$
507105143670906624	Stock 2	0.976	8.500	-0.06 ± 0.03		$0.63^{+0.07}_{-0.06}$	$0.234^{+0.043}_{-0.031}$
507119265523387136	Stock 2	0.995	8.500	-0.06 ± 0.03		–	–
507555904779576064	Stock 2	0.977	8.500	-0.06 ± 0.03		$0.35^{+0.05}_{-0.05}$	$0.118^{+0.018}_{-0.016}$
506862078583709056	Stock 2	0.999	8.500	-0.06 ± 0.03		$0.86^{+0.07}_{-0.08}$	$0.041^{+0.017}_{-0.013}$
458778927573447168	Stock 2	0.997	8.500	-0.06 ± 0.03		$0.48^{+0.09}_{-0.09}$	$0.069^{+0.021}_{-0.018}$
507362012775415552	Stock 2	0.990	8.500	-0.06 ± 0.03		$0.5^{+0.07}_{-0.07}$	$0.153^{+0.027}_{-0.029}$
507414067782288896	Stock 2	0.984	8.500	-0.06 ± 0.03		$0.29^{+0.02}_{-0.02}$	$0.028^{+0.004}_{-0.005}$
458066409683198336	Stock 2	0.994	8.500	-0.06 ± 0.03		$0.41^{+0.07}_{-0.05}$	$0.098^{+0.022}_{-0.014}$
463937282075547648	Stock 2	0.994	8.500	-0.06 ± 0.03		$0.36^{+0.05}_{-0.04}$	$0.065^{+0.012}_{-0.01}$
507128332197081344	Stock 2	0.861	8.500	-0.06 ± 0.03		$0.36^{+0.06}_{-0.05}$	$0.278^{+0.036}_{-0.03}$
507277870080186624	Stock 2	0.899	8.500	-0.06 ± 0.03		–	–
506864793008901632	Stock 2	0.698	8.500	-0.06 ± 0.03		$0.3^{+0.12}_{-0.08}$	$0.284^{+0.052}_{-0.054}$
507221863701989248	Stock 2	0.887	8.500	-0.06 ± 0.03		–	–
1992469104239732096	Stock 12	0.999	8.450			$0.35^{+0.44}_{-0.15}$	$0.127^{+0.176}_{-0.096}$

Notes. P_{WD} is the probability of the object being a WD, adopted from [Gentile Fusillo et al. \(2019\)](#), $\log t_{\text{cl}}$ is the cluster age, and $[\text{Fe}/\text{H}]/[\text{M}/\text{H}]$ is the cluster metallicity. Assuming that all recovered WDs are of the DA type, M_{WD} and t_{cool} are WD mass and WD cooling age estimates, respectively. ^(a)Recovered in [Holt et al. \(2019\)](#) but not characterized. Missing values of M_{WD} and t_{cool} for some objects are due to GDR2 photometry problems for these objects.

References. OC metallicities: [Bagdonas et al. \(2018\)](#), [Baratella et al. \(2020\)](#), [Carrera et al. \(2019\)](#), [Conrad et al. \(2014\)](#), [Fritzewski et al. \(2019\)](#), [Netopil et al. \(2016\)](#), [Netopil \(2017\)](#), [Reddy & Lambert \(2019\)](#), [Zhang et al. \(2019\)](#).

M_{WD} and t_{cool} estimates and their errors for the novel or newly characterized WDs are listed in Table 2, where the listed values correspond to the median values obtained from the simulations and the quoted errors are derived from the 68% confidence intervals.

It is apparent that we recovered mostly intermediate- and low-mass WD members. This is understandable when the properties of massive ($\gtrsim 0.9 M_{\odot}$) WDs and the magnitude limit of *Gaia* are considered. The highest-mass WDs are less luminous and cool more rapidly than their lower-mass counterparts. Thus, they remain bright enough for *Gaia* only in the

closest and youngest OCs. Additionally, high-mass WDs can be ejected from their parent OC due to the potential velocity kicks imparted on them during their formation by asymmetric mass-loss or dynamical interactions with other OC stars ([Fellhauer et al. 2003](#); [Tremblay et al. 2012](#)). Last, the number of young OCs potentially hosting sufficiently bright massive WDs in the solar neighborhood is low. Therefore, also taking the degradation of the astrometry and photometry quality of *Gaia* when approaching its magnitude limit into consideration, only very few massive WDs are recovered by our approach, as expected.

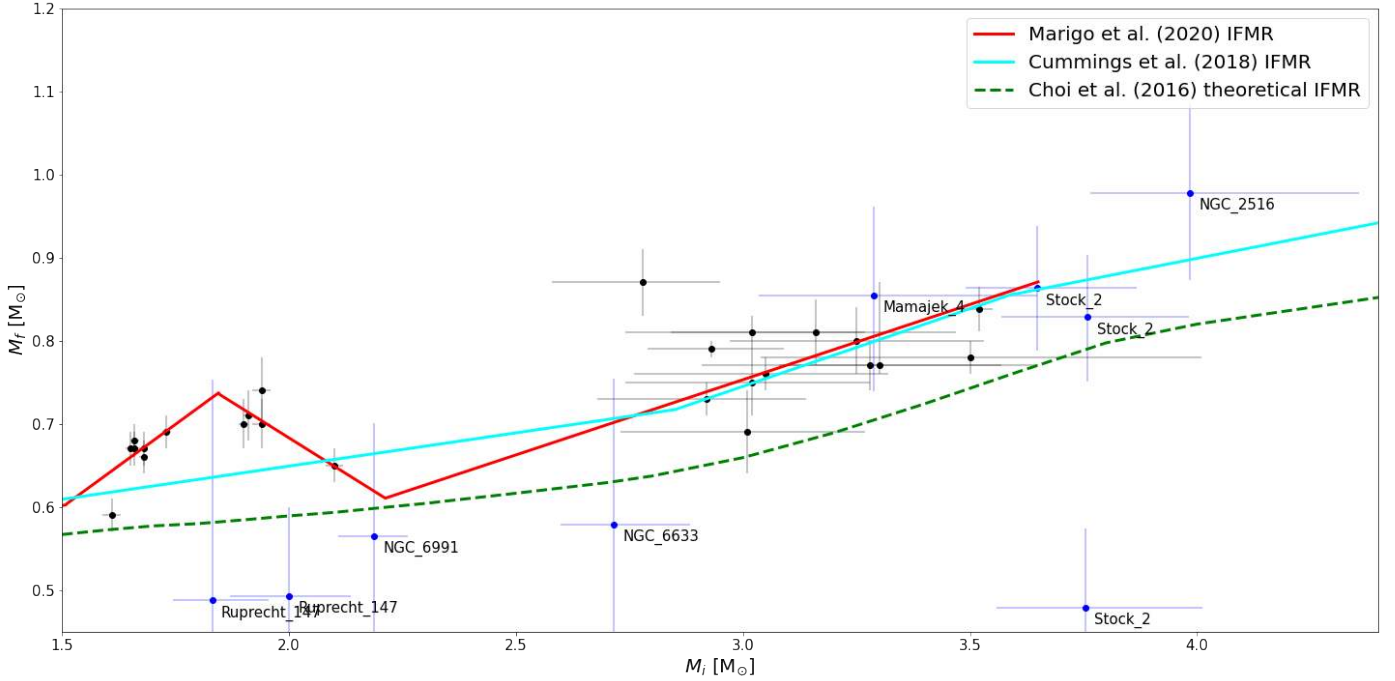


Fig. 6. Semi-empirical IFMRs in the range of M_i from 1.5 to $4.4 M_{\odot}$. The data points include the newly recovered and characterized WD OC members (in blue, with parent OC labeled, Table 4) and the previously published OC WDs from Table 3 and Marigo et al. (2020) (in black). The four-piece IFMR fit (red) is adopted from Marigo et al. (2020). The cyan line represents the IFMR fit adopted from Cummings et al. (2018), and the dashed green line is the theoretical IFMR derived from Choi et al. (2016).

Table 3. Recovered WD-OC associations previously discussed in the literature.

GDR2 source id	Associated cluster	Refs.
66697547870378368	Melotte 22	Eggen & Greenstein (1965)
3029912407273360512	NGC 2422	Richer et al. (2019)
5289447182180342016	NGC 2516	Holt et al. (2019)
5294015515555860608	NGC 2516	Holt et al. (2019)
5290767695648992128	NGC 2516	Reimers & Koester (1982)
659494049367276544	NGC 2632	Salaris & Bedin (2019)
661841163095377024	NGC 2632	Salaris & Bedin (2019)
665139697978259200	NGC 2632	Salaris & Bedin (2019)
664325543977630464	NGC 2632	Salaris & Bedin (2019)
662798086105290112	NGC 2632	Salaris & Bedin (2019)
661297901272035456	NGC 2632	Salaris & Bedin (2019)
661353224747229184	NGC 2632	Salaris & Bedin (2019)
662998983199228032	NGC 2632	Salaris & Bedin (2019)
661270898815358720	NGC 2632	Salaris & Bedin (2019)
661010005319096192	NGC 2632	Salaris & Bedin (2019)
660178942032517760	NGC 2632	Salaris & Bedin (2019)
661311267210542080	NGC 2632	Salaris & Bedin (2019)
5338718261060841472	NGC 3532	Koester & Reimers (1993)
4477214475044842368	NGC 6633	Reimers & Koester (1994)
2170776080281869056	NGC 7092	Caiazzo et al. (2020)
4088108859141437056	Ruprecht 147	Marigo et al. (2020)
4087806832745520128	Ruprecht 147	Marigo et al. (2020)
4183919237232621056	Ruprecht 147	Marigo et al. (2020)
4184169822810795648	Ruprecht 147	Marigo et al. (2020)

6. IFMR

Using the previously obtained M_{WD} and t_{cool} values and supplementing them with the values obtained from the literature, we can investigate the IFMR. In the IFMR analysis, an accurate

determination of the OC age is critical. This is particularly true for young OCs with young WDs, where the derived masses of the WD progenitors are very sensitive to the evolutionary time, which is derived from the OC age and WD cooling age.

We are interested in objects that have undergone single-star evolution, so we restricted the analysis to objects with $M_{\text{WD}} > 0.45 M_{\odot}$. Below this mass boundary, all objects are thought to be the product of close binary evolution (Tremblay et al. 2016).

If the cluster age t_{cl} and the WD cooling age t_{cool} are known, the lifetime of the progenitor can be given by $t_{\text{prog}} = t_{\text{cl}} - t_{\text{cool}}$. To calculate the progenitor mass from t_{prog} , an approximate mass-luminosity relation is commonly used for back-of-the-envelope calculations:

$$L/L_{\odot} \sim (M/M_{\odot})^{\alpha}. \quad (5)$$

In order to obtain more credible results, we used PARSEC version 1.2S (Bressan et al. 2012) and COLIBRI S_35 (Pastorelli et al. 2019) isochrones⁴ to determine the initial mass of the progenitor. For each WD, we performed 100 Monte Carlo simulations, each time drawing a value from the normal distribution of t_{cl} , cluster metallicity, and t_{cool} distribution obtained in the previous section. All distributions were assumed to be independent. Since t_{cl} measurements generally lack uncertainties, the 1σ error for t_{cl} was assumed to be 10% of its measured value. The metallicity distribution was also centered on its measured value, with 1σ being its uncertainty as adopted from the literature. The initial progenitor masses M_i and their errors were obtained in the same way as M_{WD} and t_{cool} in the previous section. The resulting IFMR is plotted in Fig. 6.

It can be seen from Fig. 6 that the newly characterized WDs are consistent with the nonlinear IFMR from Marigo et al. (2020), with a kink located over $1.65 M_{\odot} \lesssim M_i \lesssim 2.1 M_{\odot}$, which

⁴ http://stev.oapd.inaf.it/cgi-bin/cmd_3.3

Table 4. Initial progenitor masses M_i for the newly characterized WDs in Fig. 6.

GDR2 source id	Associated cluster	M_i (M_\odot)
6653447981289591808	Mamajek 4	$3.3^{+0.4}_{-0.3}$
5294015515555860608	NGC 2516	$4.0^{+0.4}_{-0.2}$
4476643725433841920	NGC 6633	$2.7^{+0.2}_{-0.1}$
2166915179559503232	NGC 6991	$2.2^{+0.1}_{-0.1}$
4183928888026931328	Ruprecht 147	$1.8^{+0.1}_{-0.1}$
4183926006112672768	Ruprecht 147	$2.0^{+0.1}_{-0.1}$
507054806657042944	Stock 2	$3.8^{+0.3}_{-0.2}$
506862078583709056	Stock 2	$3.6^{+0.2}_{-0.2}$
458778927573447168	Stock 2	$3.8^{+0.3}_{-0.2}$

they interpreted as a signature of the lowest-mass stars in the Galaxy that become carbon stars during the thermally pulsing asymptotic giant branch phase. Of particular interest are the WDs hosted by NGC 6991 and NGC 6633, which fall into the IFMR dip that, until then, had not been sufficiently characterized. There are also other WDs that fall into this gap (members of IC 4756, Alessi 62, and NGC 2527), which were either below the mass cutoff or had problems in their GDR2 parameters. The three-piece IFMR fit from Cummings et al. (2018) and the theoretical IFMR adopted from Choi et al. (2016) are also shown in Fig. 6. It can be seen that the IFMR fits of Marigo et al. (2020) and Cummings et al. (2018) are almost identical from $M_i \gtrsim 2.9 M_\odot$.

Apart from the IFMR kink at $1.65 M_\odot \lesssim M_i \lesssim 2.1 M_\odot$, there is a visible offset between the theoretical and observed masses from approximately $M_i \gtrsim 3.0 M_\odot$, where the observed WD masses are $\sim 0.1 M_\odot$ more massive than predicted, as has been noted in Cummings et al. (2018). Cummings et al. (2019) have later attributed this offset mainly to the effects of convective-core overshoot and rotational mixing in the main-sequence progenitors, where the rotational effects are not taken into consideration in the theoretical IFMR models. The newly characterized OC WDs with $M_i \gtrsim 3.0 M_\odot$ also continue to follow this trend, being $\sim 0.1 M_\odot$ more massive than what the theoretical IFMRs (e.g., Choi et al. 2016) predict.

Other WDs below the IFMR fit are most likely binaries, or possibly foreground objects, that have been incorrectly assigned to the OC. Interestingly, Stock 2 seems to host a large number of WDs scattered in the IFMR; some of them follow the IFMR fit by Marigo et al. (2020), but others are clustered around $M_{\text{WD}} = 0.4 M_\odot$. Such WDs may be members of binary systems. Additional scatter can be attributed to the effects of strong and variable extinction, which has been noted for this cluster (Spagna et al. 2009).

White dwarfs are the final products of the evolution of stars with initial masses (assuming solar metallicity) less than $8\text{--}10 M_\odot$ (Langer 2012; Smartt 2009); however, in binary systems, the initial mass for one of the components can be as high as $15 M_\odot$ (Wellstein et al. 2001) or as low as $6 M_\odot$ (Podsiadlowski et al. 2004). Finding a high-mass WD in a young OC can help identify initial masses for stars that undergo electron-capture SNe. We managed to identify one potential high-mass WD in NGC 2516. However, its cooling time only suggests a $\sim 4 M_\odot$ progenitor. Due to the shortcomings of this analysis, as described above, we did not recover any other high-mass

WDs and are therefore unable to put any new constraints on the boundary between neutron stars and WD formation.

7. Summary and conclusions

We searched for new potential WDs that are possible OC members using the WD catalog by Gentile Fusillo et al. (2019) and the OC catalog by Cantat-Gaudin et al. (2018a), both based on GDR2 data. Such associations are very valuable as ascertaining the membership of a WD to an OC allows us to adopt the OC distance to the WD. This distance is more precise than the distance determined from the WD parallax by itself as it is based on a large number of stars and because the WD parallaxes in the GDR2 exhibit high uncertainties due to their faintness and blue colors. This enables a more precise determination of the WD parameters. Furthermore, the nature of OCs as a coeval group of stars with a common origin allows us to study a number of topics, such as IFMR and metallicity effects.

Our study confirmed the cluster membership of several literature WD cluster members and uncovered a number of new associations. On the other hand, there are a lot of established literature OC WDs that do not seem to satisfy the astrometric and photometric criteria for cluster membership in the GDR2. Removing them from IFMR studies may alleviate the scatter that is present in the data.

The derived WD and progenitor masses of the novel WDs are broadly in line with the IFMR fit of Marigo et al. (2020), although a large number of binaries falling below the fit are also likely present. Some of the recovered WDs from NGC 6991 and NGC 6633 fall into the IFMR dip, which has been poorly characterized and deserves further study. There are several WDs lying in this gap that had to be discarded from the analysis due to their low derived masses (possibly due to binarity with a low-mass companion) or problems with the GDR2 photometry or astrometric solution (such as WDs hosted by IC 4756, Alessi 62, and NGC 2527). It could be worthwhile to observe these objects spectroscopically or revisit them in the next *Gaia* data release.

This work showcases the possibilities that precise astrometry can bring to WD studies. Naturally, spectroscopic observations of the WD cluster member candidates are still needed to confirm their WD status and type, as well as to provide more precise parameters and an additional check for cluster membership.

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Appendix A: Proper motion diagrams, parallax distributions, and CMDs of the OC-WD associations

from Sect. 3. They are either novel candidates or were gathered from the literature.

In this section, we provide the proper motion diagrams, parallax distributions, and CMDs for the rest of the cluster-WD pairs

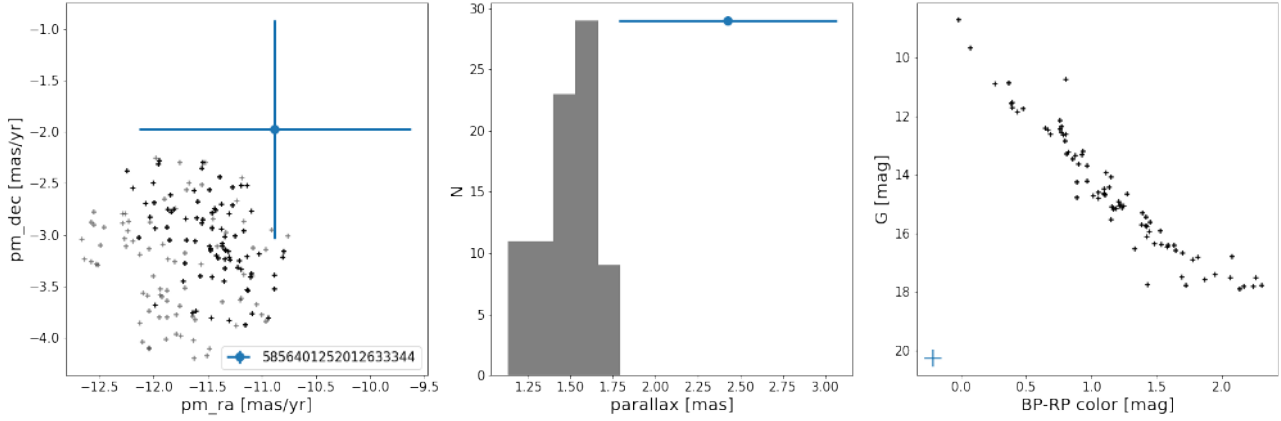


Fig. A.1. Same as in Fig. 5, but for ASCC 73.

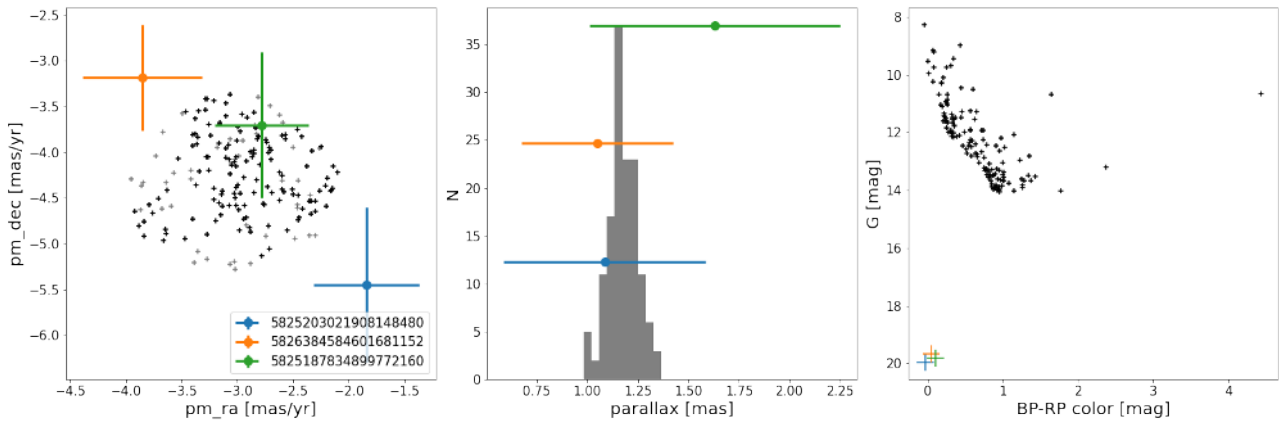


Fig. A.2. Same as in Fig. 5, but for ASCC 79.

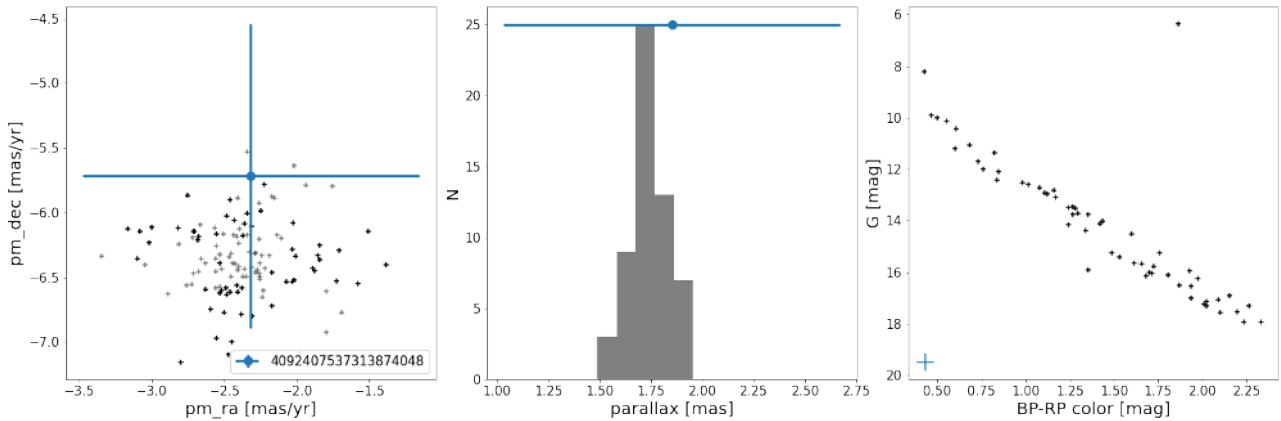


Fig. A.3. Same as in Fig. 5, but for ASCC 97.

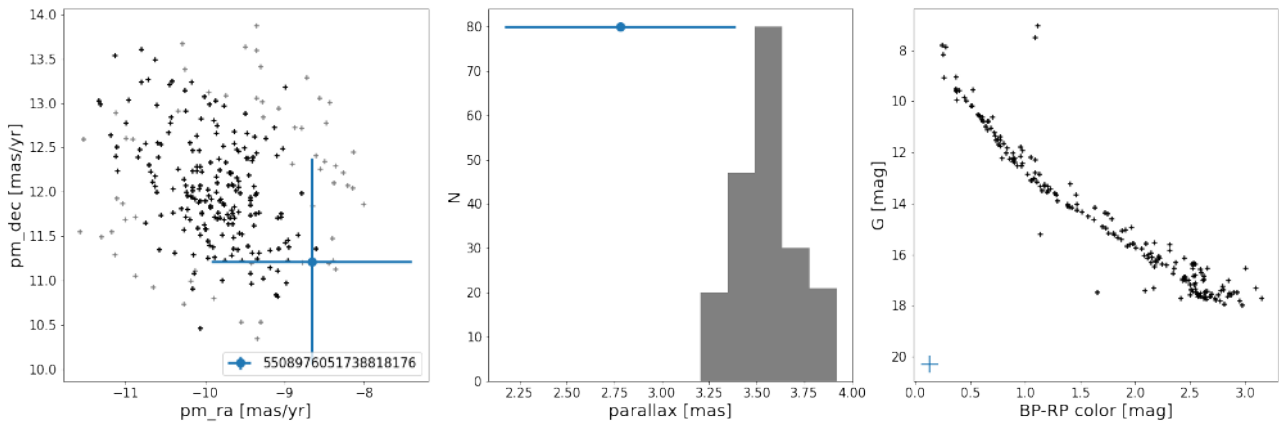


Fig. A.4. Same as in Fig. 5, but for Alessi 3.

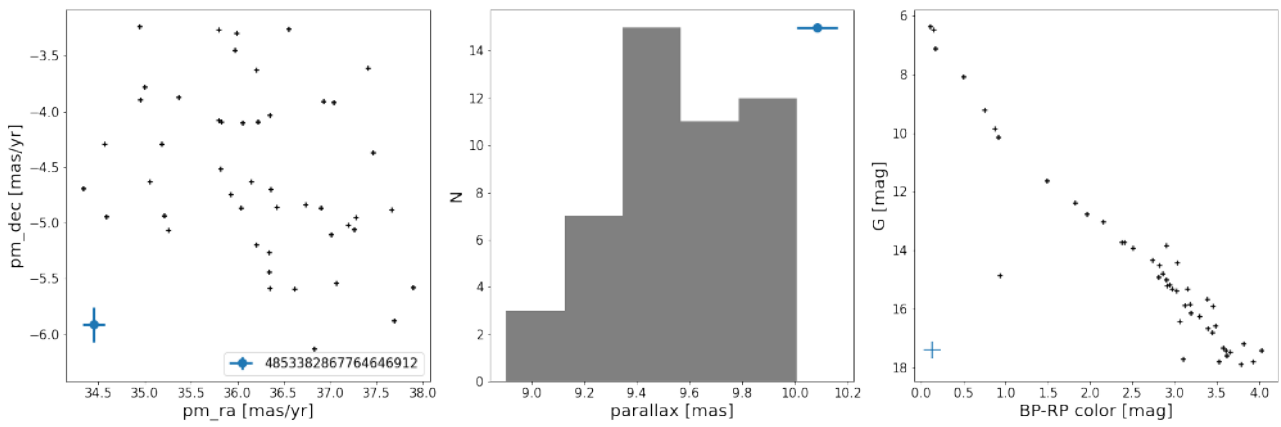


Fig. A.5. Same as in Fig. 5, but for Alessi 13.

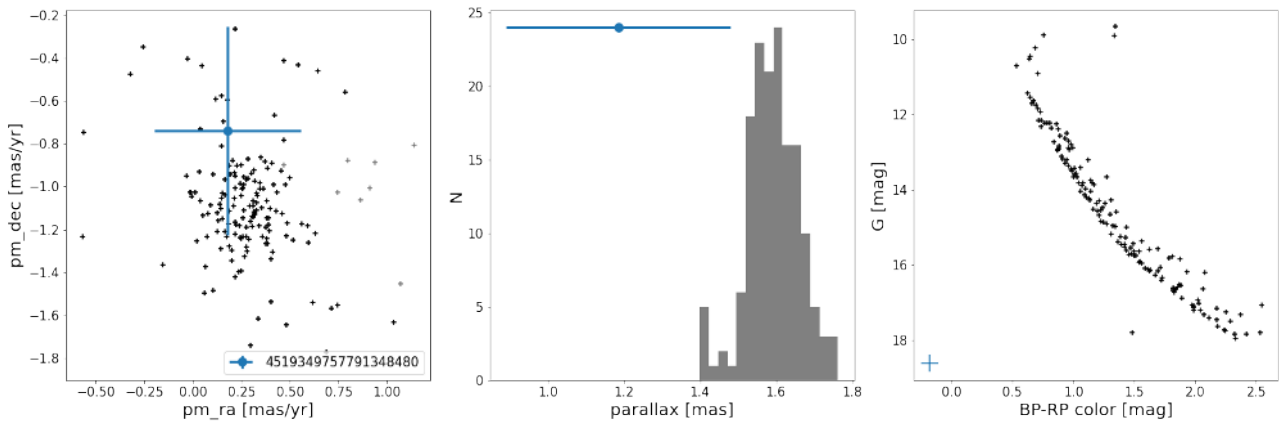


Fig. A.6. Same as in Fig. 5, but for Alessi 62.

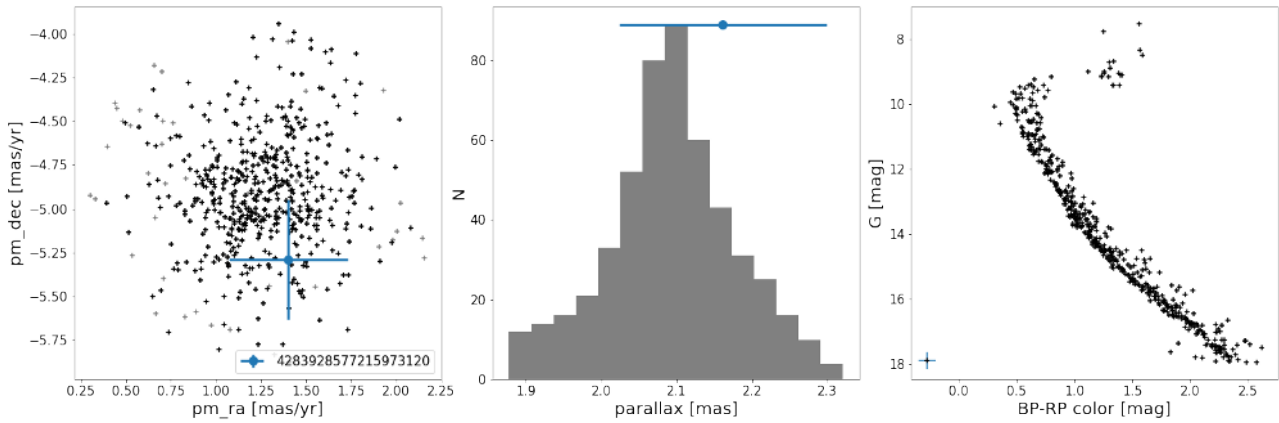


Fig. A.7. Same as in Fig. 5, but for IC 4756.

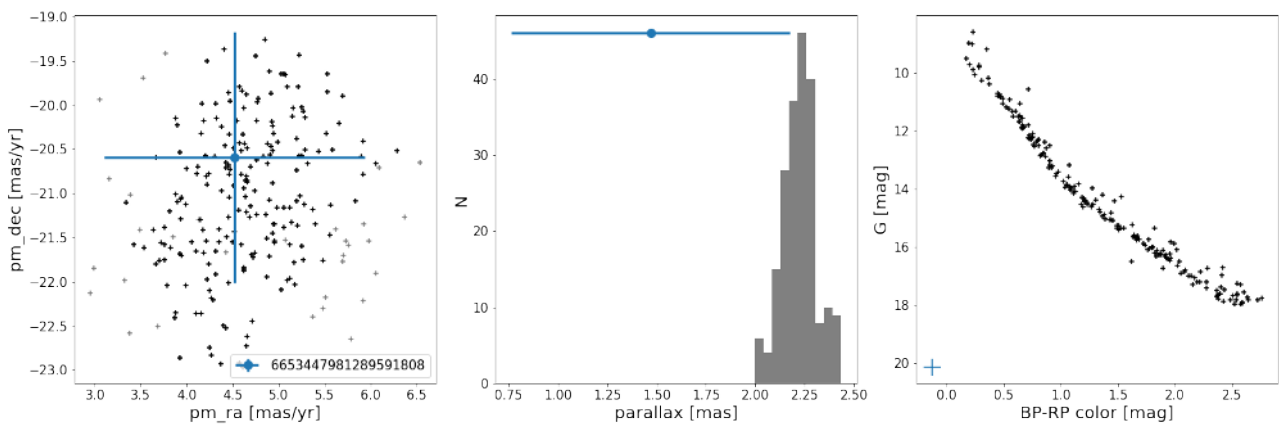


Fig. A.8. Same as in Fig. 5, but for Mamajek 4.

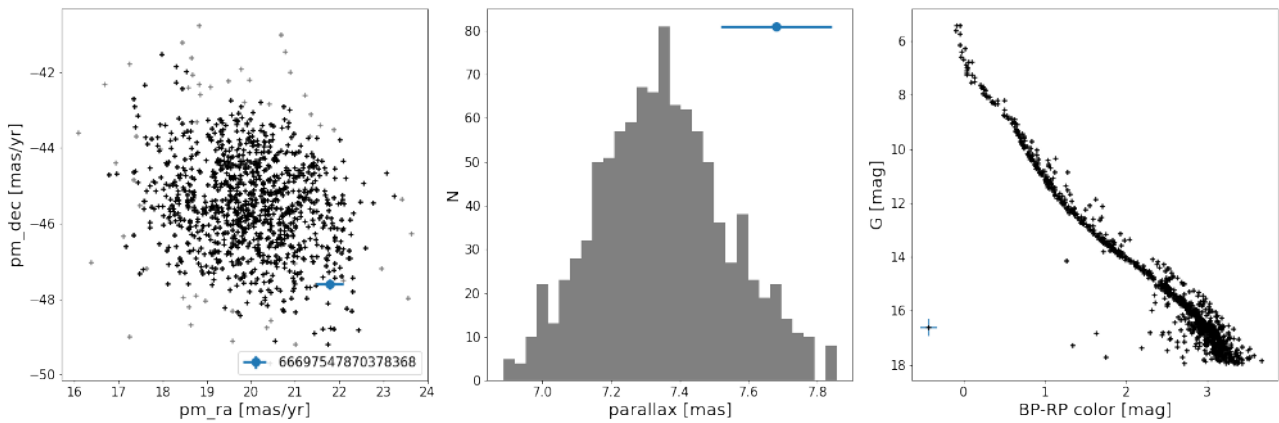


Fig. A.9. Same as in Fig. 5, but for Melotte 22.

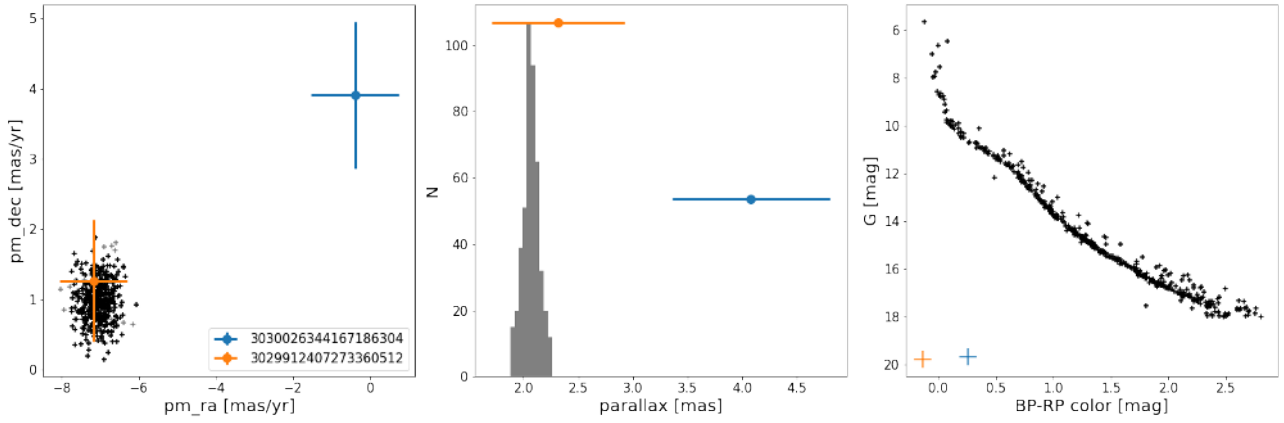


Fig. A.10. Same as in Fig. 5, but for NGC 2422.

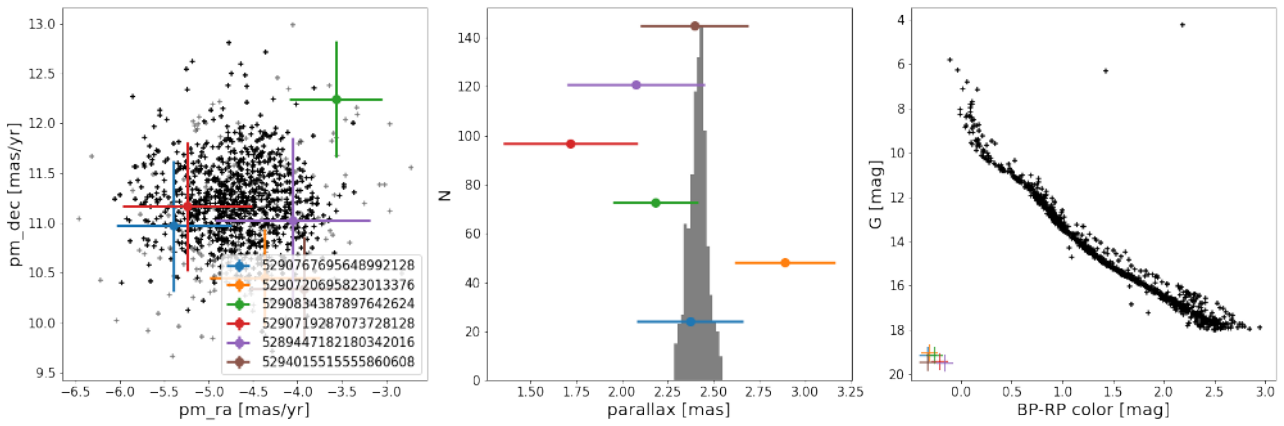


Fig. A.11. Same as in Fig. 5, but for NGC 2516.

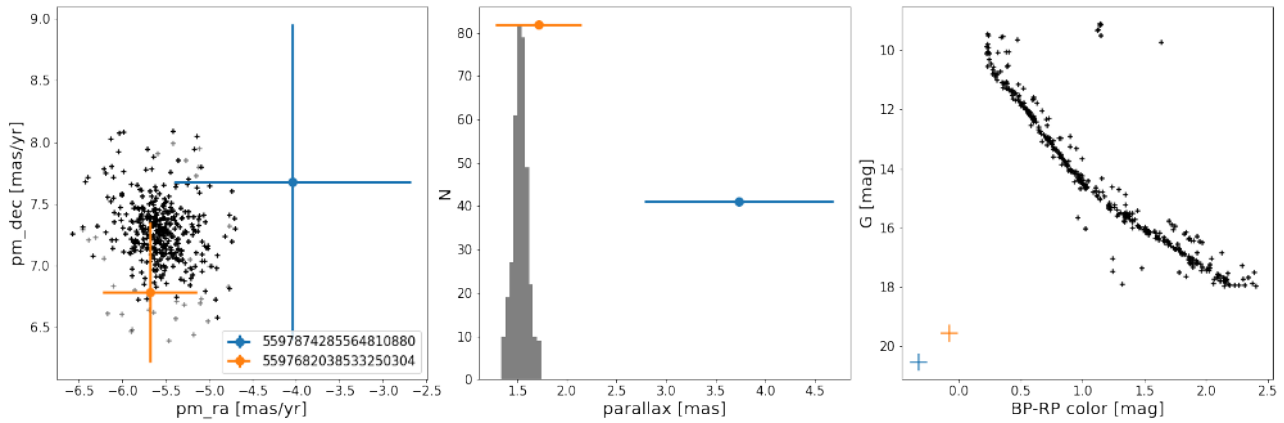


Fig. A.12. Same as in Fig. 5, but for NGC 2527.

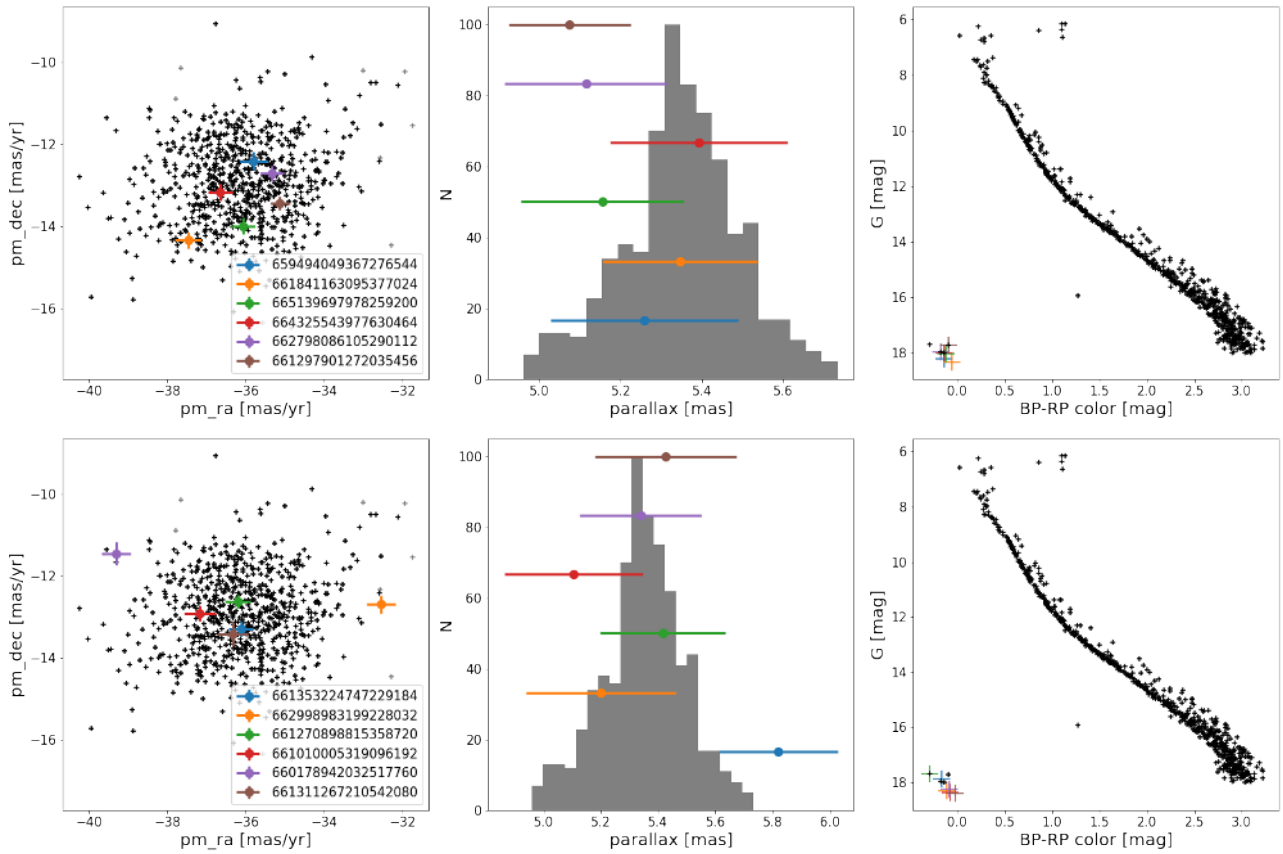


Fig. A.13. Same as in Fig. 5, but for NGC 2632.

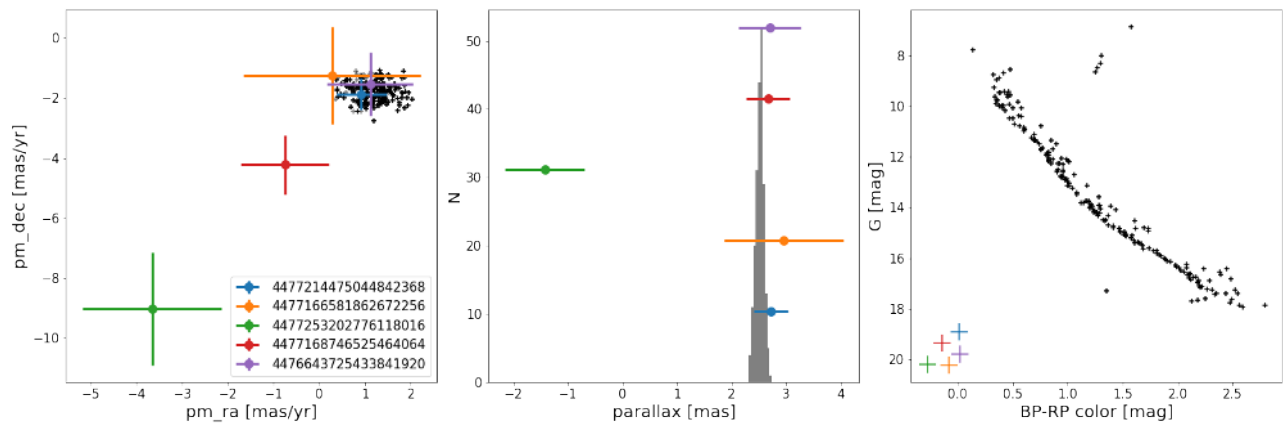


Fig. A.14. Same as in Fig. 5, but for NGC 6633.

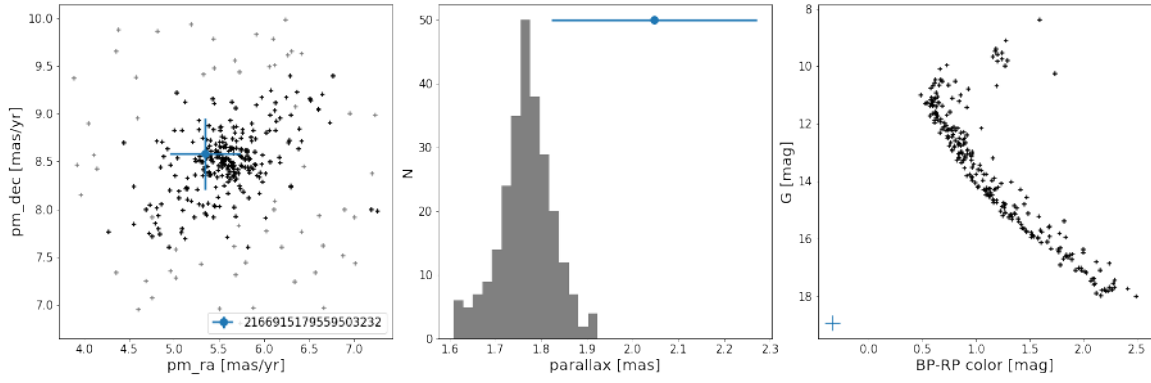


Fig. A.15. Same as in Fig. 5, but for NGC 6991.

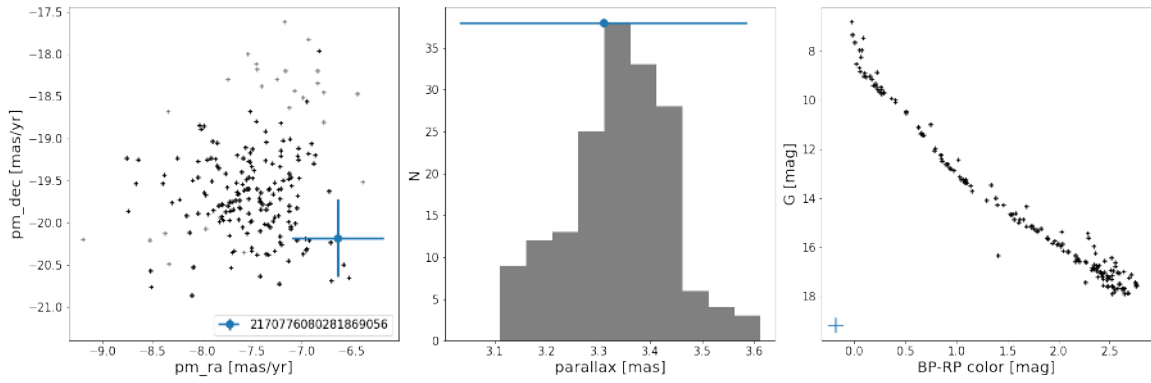


Fig. A.16. Same as in Fig. 5, but for NGC 7092.

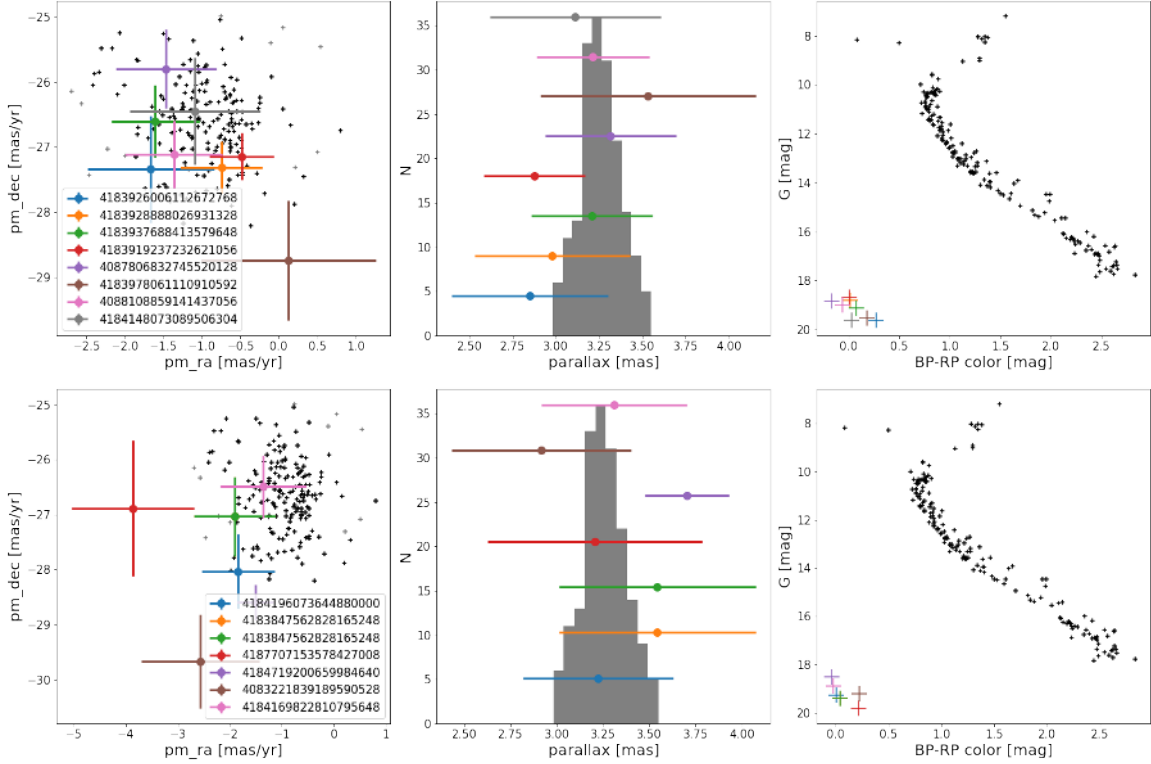


Fig. A.17. Same as in Fig. 5, but for Ruprecht 147.

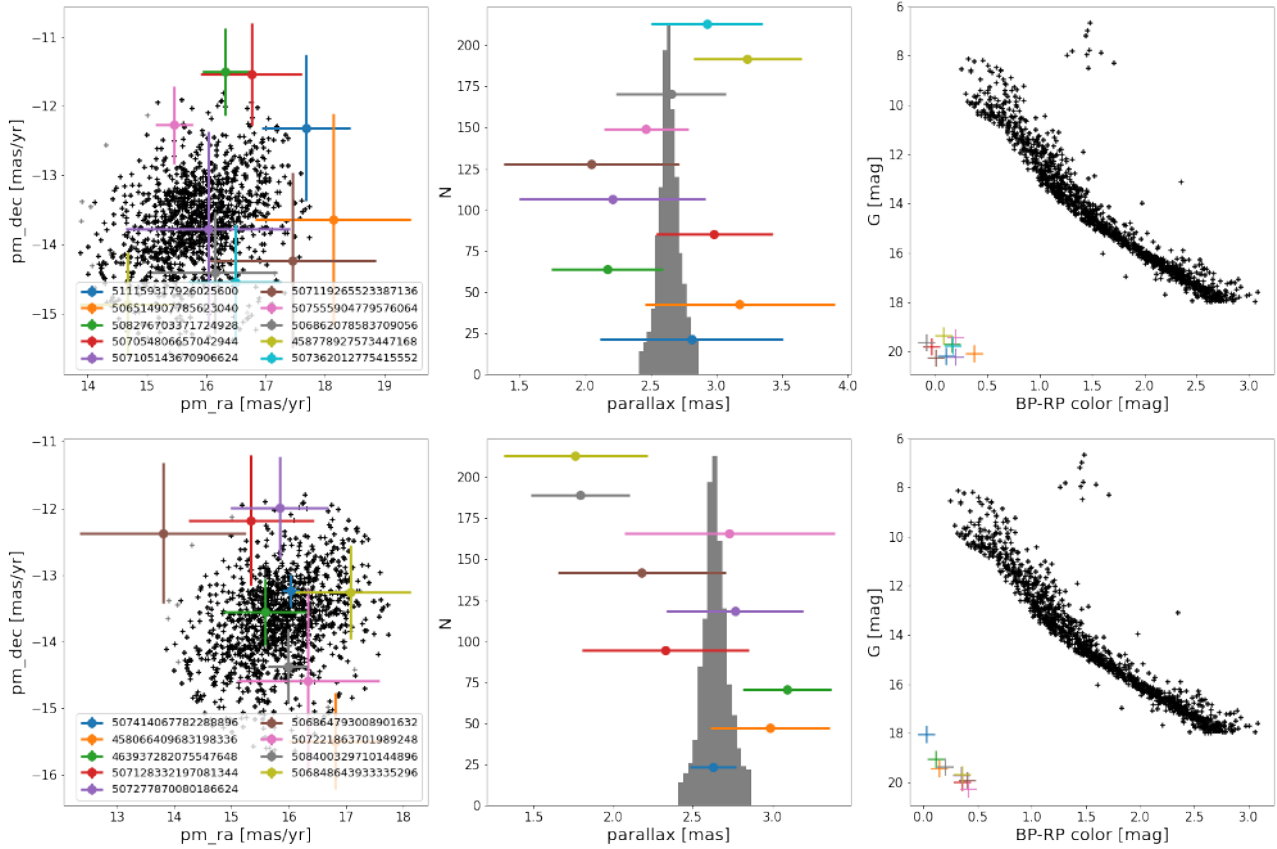


Fig. A.18. Same as in Fig. 5, but for Stock 2.

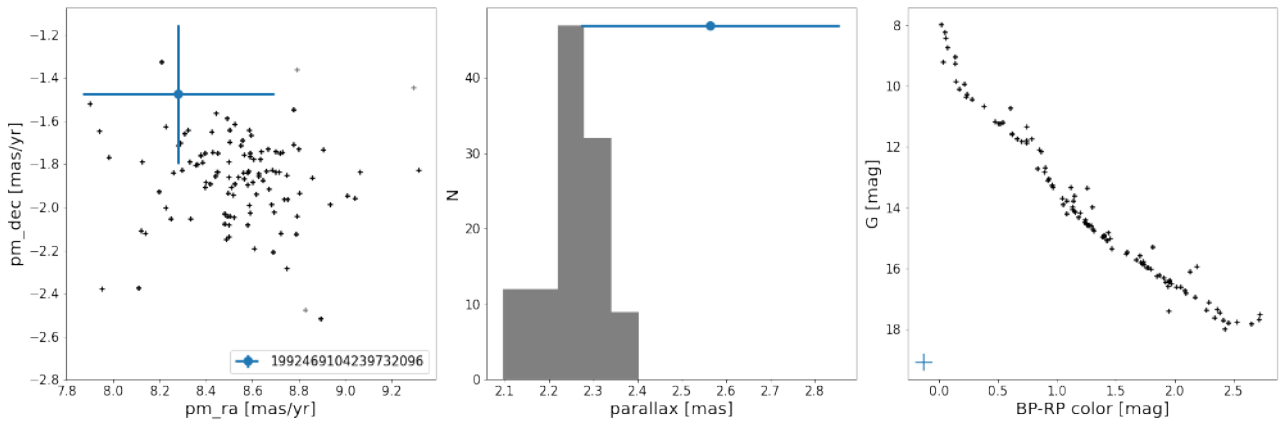


Fig. A.19. Same as in Fig. 5, but for Stock 12.