



## WISEA J114724.10-204021.3: A FREE-FLOATING PLANETARY MASS MEMBER OF THE TW HYA ASSOCIATION

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### ABSTRACT

We present WISEA J114724.10-204021.3, a young, low-mass, high-probability member of the TW Hya association (TWA). WISEA J114724.10-204021.3 was discovered based on its red AllWISE color ( $W1 - W2 = 0.63$  mag) and extremely red 2MASS  $J - K_S$  color ( $>2.64$  mag), the latter of which is confirmed with near-infrared photometry from the Visible and Infrared Survey Telescope for Astronomy Hemisphere Survey ( $J - K_S = 2.57 \pm 0.03$ ). Follow-up near-infrared spectroscopy shows a spectral type of  $L7 \pm 1$  as well as several spectroscopic indicators of youth. These include a peaked  $H$ -band shape and a steeper  $K$ -band slope, traits typically attributed to low surface gravity. The sky position, proper motion, and distance estimates of WISEA J114724.10-204021.3 are all consistent with membership in the  $\sim 10$  Myr old TWA. Using the age of the TWA and evolutionary models, we estimate the mass of WISEA J114724.10-204021.3 to be 5–13  $M_{\text{Jup}}$ , making it one of the youngest and lowest-mass free-floating objects yet discovered in the Solar neighborhood.

*Key words:* brown dwarfs

*Supporting material:* data behind figure

### 1. INTRODUCTION

Young, late-type L dwarfs (brown dwarfs with ages  $<100$  Myr and  $T_{\text{eff}} \lesssim 1600$  K) show striking spectroscopic similarities to young, directly imaged exoplanets (e.g., Gizis et al. 2012; Faherty et al. 2013; Liu et al. 2013). Unlike exoplanets, however, free-floating brown dwarfs are much simpler to observe because they lack a nearby, bright host Sun. As such, young, low-mass brown dwarfs are ideal laboratories for investigating the physical conditions likely to be present in giant exoplanets, thus offering critical checks of theory. While young brown dwarfs belonging to nearby ( $d < 100$  pc) young associations are beginning to be found in greater numbers (e.g., Gagné et al. 2015), very few young, late-type L ( $>L5$ ) dwarfs are currently known ( $\lesssim 10$ ; see Figure 4 of Gagné et al. 2015). However, it is these young, late-type L dwarfs that are crucial as comparison objects for directly imaged young brown dwarf and planetary companions (e.g., Bowler et al. 2014; Bonnefoy et al. 2016; Gauza et al. 2015; Hinkley et al. 2015).

If a substellar object can be connected to a young, nearby moving group, then its age can be firmly established, thereby providing a vital anchor point for low-mass evolutionary models. The young ( $10 \pm 3$  Myr; Bell et al. 2015) TW Hya association (TWA) is one of the nearest regions of recent star formation. Its proximity ( $\sim 50$  pc) and young age make it an excellent testbed for studying early phases of stellar and substellar evolution. Here, we report the discovery of WISEA J114724.10-204021.3, a new L7 member of the TWA.

### 2. IDENTIFICATION OF WISEA J114724.10-204021.3

As noted in Schneider et al. (2014), young, late-type L dwarfs occupy a unique region of 2MASS (Skrutskie

et al. 2006) and *Wide-field Infrared Survey Explorer* (WISE; Wright et al. 2010) color space compared to field L dwarfs because of their extremely red near-infrared colors (see their Figure 5). WISEA J114724.10-204021.3 (hereafter WISEA 1147–2040) was found as part of a larger program focused on finding young, late-type L dwarfs based on their 2MASS and AllWISE<sup>5</sup> colors. Briefly, candidate young L dwarfs were chosen by requiring that they have a 2MASS  $J - K_S$  color between 2.0 and 3.5 mag and an AllWISE  $W1$  ( $3.6 \mu\text{m}$ ) –  $W2$  ( $4.5 \mu\text{m}$ ) color between 0.3 and 0.9 mag. Candidates were also required to have their  $K_S$ ,  $W1$ , and  $W2$  passband uncertainties be greater than zero (i.e., not upper limits). Note that we did not require objects to be well detected in the  $J$  band, anticipating the existence of objects so red in  $J - K_S$  color that they may be detected at the  $K_S$  band, but not at  $J$  (as is the case for WISEA 1147–2040). Last, we required that the separation between the 2MASS and AllWISE source positions of candidates be greater than  $1''$ , thereby ensuring each candidate shows appreciable proper motion between the 2MASS and AllWISE epochs ( $\gtrsim 100$  mas  $\text{yr}^{-1}$ , considering the  $\sim 10$  year time baseline between 2MASS and AllWISE). We then scrutinized each candidate individually by inspecting available optical (DSS and SDSS), near-infrared (2MASS), and mid-infrared (AllWISE) images to ensure each candidate is a point source (i.e., not extended or blended) with noticeable proper motion. Of the  $\sim 50$  returned candidate red L dwarfs, WISEA 1147–2040 was picked out early as a strong young L dwarf candidate because of its uncommonly red 2MASS  $J - K_S$  color ( $>2.64$  mag) and therefore worthy of follow-up spectroscopic observations. The basic properties and photometry of WISEA 1147–2040 are given in Table 1.

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<sup>5</sup> <http://http://wise2.ipac.caltech.edu/docs/release/allwise/expsup/>

**Table 1**  
WISEA J114724.10-204021.3 Properties

Parameter	Value	Ref.
Identifiers		
AllWISE	J114724.10-204021.3	1
2MASS	11472421-2040204	2
Observed Properties		
$\alpha$ (J2000)	11:47:24.10	1
$\delta$ (J2000)	-20:40:21.3	1
$\mu_\alpha$	$-122.1 \pm 12.0$ mas yr $^{-1}$	4
$\mu_\delta$	$-74.5 \pm 11.3$ mas yr $^{-1}$	4
$J$ (2MASS)	$>17.511$ mag	2
$H$ (2MASS)	$15.764 \pm 0.112$ mag	2
$K_S$ (2MASS)	$14.872 \pm 0.106$ mag	2
$Y$ (VHS)	$19.160 \pm 0.067$ mag	3
$J$ (VHS)	$17.445 \pm 0.029$ mag	3
$K_S$ (VHS)	$14.872 \pm 0.011$ mag	3
W1	$13.718 \pm 0.026$ mag	1
W2	$13.090 \pm 0.030$ mag	1
W3	$>12.155$ mag	1
W4	$>8.913$ mag	1
$J - K_S$ (2MASS)	$>2.64$ mag	2
$J - K_S$ (VHS)	$2.57 \pm 0.03$ mag	3
$J - K_S$ (synthetic)	$2.61 \pm 0.03$ mag	4
W1 - W2	$0.63 \pm 0.04$ mag	1
Inferred Properties		
Spectral Type (NIR)	$L7 \pm 1$ (red)	4
Photometric Distance	$31.2 \pm 1.5$ pc	4
Kinematic Distance	32–33 pc	4
Predicted Distance	$31.3 \pm 3.8$ pc	4, 5
$T_{\text{eff}}^a$	$1500 \pm 100$ K	4
$\log g^a$	$4.0 \pm 0.5$	4
Mass <sup>a</sup>	6–13 $M_{\text{Jup}}$	4
$T_{\text{eff}}^b$	1100–1200 K	4
Mass <sup>b</sup>	5–6 $M_{\text{Jup}}$	4
Age	$10 \pm 3$ Myr	6

#### Notes.

<sup>a</sup> From model fitting (Section 4.2).

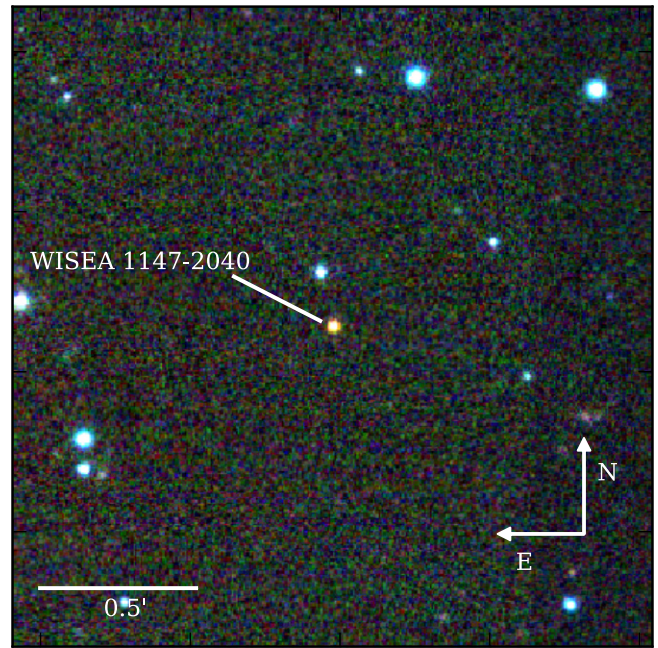
<sup>b</sup> From the estimated bolometric luminosity (Section 5).

**References.** (1) AllWISE; (2) 2MASS; (3) VHS; (4) This work; (5) BANYAN II (Malo et al. 2013; Gagné et al. 2014); (6) Bell et al. (2015).

### 3. OBSERVATIONS

#### 3.1. Infrared Telescope Facility (IRTF)/SpeX

WISEA 1147–2040 was observed with the upgraded SpeX spectrograph (Rayner et al. 2003) on the night of 2016 February 12 UT at the NASA IRTF on Mauna Kea. The observations were made in prism mode with a  $0''.5$  wide slit, which achieves a resolving power ( $\lambda/\Delta\lambda$ ) of  $\sim 150$  over the range  $0.8\text{--}2.5\ \mu\text{m}$ . We oriented the  $15''$  long slit along the parallactic angle and obtained a series of eighteen 120 s exposures at two different nod positions along the slit for a total exposure time of 2160 s. The A0V star HD 101122 was observed at a similar airmass for telluric correction purposes. The spectrum was reduced using the SpeXtool reduction package (Vacca et al. 2003; Cushing et al. 2004). The extracted spectrum was flux calibrated using the VHS  $K_S$  photometry.



**Figure 1.** Three-color composite VHS image centered on the position of WISEA 1147–2040 ( $Y$ —blue,  $J$ —green,  $K_S$ —red).

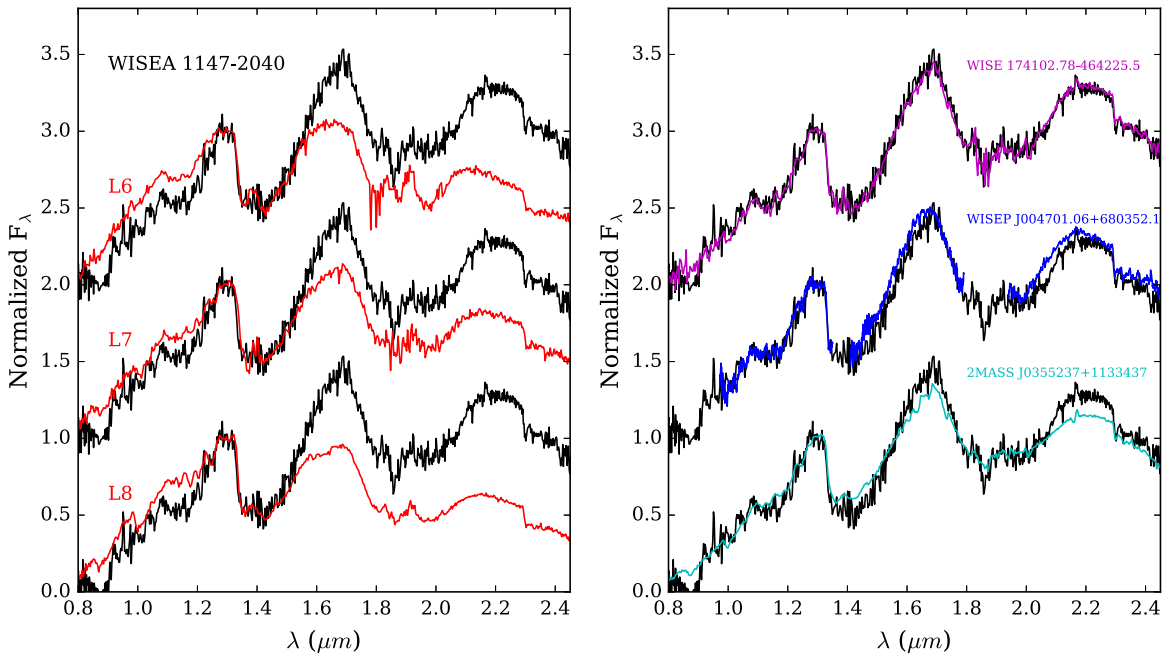
#### 3.2. Visible and Infrared Survey Telescope for Astronomy (VISTA)/VHS

WISEA 1147–2040 was observed by VISTA (Emerson et al. 2004) in the  $Y$ ,  $J$ , and  $K_S$  photometric bands (Figure 1) as part of the VISTA Hemisphere Survey (VHS; PI: McMahon, Cambridge, UK). WISEA 1147–2040 was well detected in all three bands, and the photometry from VHS is listed in Table 1. These observations confirm the very red  $J - K_S$  color of WISEA 1147–2040 ( $J - K_S = 2.57 \pm 0.03$  mag).

### 4. ANALYSIS

#### 4.1. Spectral Typing

We determined a spectral type for WISEA 1147–2040 following the method outlined in the appendix of Schneider et al. (2014). A comparison of the SpeX spectrum of WISEA 1147–2040 with several near-infrared spectral standards (Kirkpatrick et al. 2010) from the SpeX Prism Spectral Library (Burgasser 2014) is shown in the left panel of Figure 2. While none of the L-type standards provide a good match to the spectrum of WISEA 1147–2040, the  $J$ -band portion most closely resembles that of the L7 standard. We also compare the spectrum of WISEA 1147–2040 to the known, young, late-type L dwarfs 2MASS J03552337+1133437 (Reid et al. 2006; Faherty et al. 2013), WISEP J004701.06+680352.1 (Gizis et al. 2012), and WISE J174102.78–464225.5 (Schneider et al. 2014) in the right panel of Figure 2. WISEA 1147–2040 best matches the spectrum of WISE J174102.78–464225.5, an L7 type brown dwarf with likely membership in the  $\beta$  Pictoris or AB Doradus moving groups (Schneider et al. 2014). Note that Gagné et al. (2015) type WISE J174102.78–464225.5 as L5:L7: $\gamma$ . Based on these comparisons, we estimate a spectral type of  $L7 \pm 1$  (very red) for WISEA 1147–2040.



**Figure 2.** Left: the IRTF/SpeX spectrum of WISEA 1147–2040 (black) compared with the near-infrared L6 (2MASSI J1010148–040649; Reid et al. 2006), L7 (2MASSI J0103320+193536; Cruz et al. 2004), and L8 (2MASSW J1632291+190441; Burgasser 2007b) standards (red). Right: the IRTF/SpeX spectrum of WISEA 1147–2040 (black) compared with 2MASS J03552337+1133437 (L5 $\gamma$ ; Faherty et al. 2013), WISEP J004701.06+680352.1 (L7 INT-G; Gizis et al. 2015), and WISE J174102.78–464225.5 (L7, very red; Schneider et al. 2014). Each spectrum is normalized by the mean flux from 1.27 to 1.32  $\mu\text{m}$ .

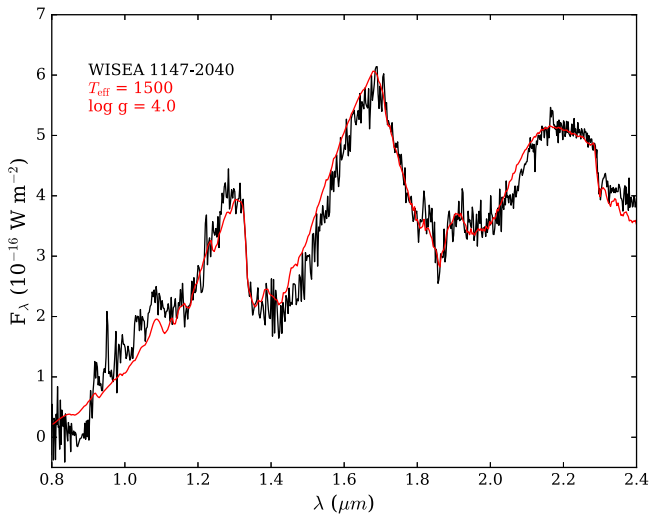
#### 4.2. Evidence of Youth

WISEA 1147–2040 has several spectroscopic and photometric traits that provide strong evidence of its young age. Because young brown dwarfs are still contracting to their final radii, they have lower surface gravities than field age brown dwarfs with the same mass. One of the consequences thought to be due to low surface gravity is an unusually dusty atmosphere. Such excessively dusty atmospheres then give rise to very red near-infrared colors. WISEA 1147–2040 has an extremely red  $J - K_S$  color ( $2.57 \pm 0.03$  mag) determined from its VHS photometry. In fact, only three free-floating L dwarfs are known to have redder  $J - K_S$  colors; the 20 Myr old  $\beta$  Pictoris moving group member PSO J318.5338–22.8603 ( $J - K_S = 2.837$  mag; Liu et al. 2013), the extremely dusty L7 dwarf ULAS J222711–004547 ( $J - K_S = 3.04$  mag; Marocco et al. 2014), and the recently discovered TWA candidate member 2MASS J11193254–1137466 ( $J - K_S = 2.62$  mag; Kellogg et al. 2015). While there are a few examples of L dwarfs with very red near-infrared colors that are not young (e.g., Kirkpatrick et al. 2010), bona fide young L dwarfs are typically found to have redder near-infrared colors than field dwarfs of the same spectral type (Cruz et al. 2009; Faherty et al. 2013). WISEA 1147–2040’s red  $J - K_S$  color is the first indication that it is young.

The near-infrared spectrum of WISEA 1147–2040 has a distinctly peaked  $H$ -band appearance and a steeper  $K$ -band slope. For a field age brown dwarf with a normal surface gravity, the  $H$ - and  $K$ -band portions of their spectra are predominantly shaped by  $\text{H}_2\text{O}$  and collisionally induced absorption (CIA) of  $\text{H}_2$ . For a young brown dwarf, where the surface gravity is much lower, the effects of CIA are greatly reduced, resulting in the triangular  $H$ -band and steeper  $K$ -band shapes seen in their spectra (Rice et al. 2011). Allers & Liu

(2013) defined the  $H$ -cont index to assess how peaked the  $H$ -band portion of a spectrum is. We measure an  $H$ -cont index value of 0.968 for WISEA 1147–2040, which is decidedly different than the field L dwarf population, instead aligning well with other low-gravity objects (see Figure 23 of Allers & Liu 2013 and Figure 5 of Gagné et al. 2015). Canty et al. (2013) defined the  $\text{H}_2(K)$  index as a measure of the  $K$ -band continuum shape and showed it could easily distinguish young brown dwarfs from field brown dwarfs for late-M spectral types. Schneider et al. (2014) extended the  $\text{H}_2(K)$  index into the L dwarf regime and showed that it could also be used to distinguish low gravity for L spectral types. We measure an  $\text{H}_2(K)$  value for WISEA 1147–2040 of 1.035, again aligning well with other low surface gravity L dwarfs (see Figure 10 of Schneider et al. 2014 and Figure 14 of Gagné et al. 2015). Other low-gravity indices (e.g.,  $\text{FeH}_2$ ,  $\text{VO}_2$ ) are only functional for spectral types  $\lesssim \text{L5}$  and are therefore unsuitable for WISEA 1147–2040 (Allers & Liu 2013).

We can also find evidence of whether or not WISEA 1147–2040 has a low surface gravity by comparing its near-infrared spectrum to models with varying surface gravities. We compare the atmospheric models of Allard et al. (2012) using the method of Cushing et al. (2008). We find a best-fitting temperature of  $1500 \pm 100$  K and surface gravity of  $4.0 \pm 0.5$ , a surface gravity much lower than a typical field L dwarf. The best-fitting model is shown in Figure 3. The combination of an extremely red  $J - K_S$  color, triangular  $H$ -band shape, steeper  $K$ -band slope, and a low surface gravity estimate from model fitting lead us to conclude that WISEA 1147–2040 has a low surface gravity and is therefore young. Kirkpatrick et al. (2008) show that low-gravity features only manifest in objects with ages less than that of the Pleiades, so we take  $\sim 100$  Myr as the upper age limit for WISEA 1147–2040.



**Figure 3.** IRTF/SpEx spectrum of WISEA 1147–2040 (black) compared with the best-fitting BT-Settl model (red;  $T_{\text{eff}} = 1500$  K,  $\log g = 4.0$ ). The data used to create this figure are available.

### 4.3. Membership in the TW Hya Association

Because WISEA 1147–2040 shows youthful characteristics, we can now evaluate whether or not it belongs to one of the young nearby associations. We first use the BANYAN II (Malo et al. 2013; Gagné et al. 2014) moving group membership evaluation tool, which takes the position and proper motion of a source and, through the use of a naive Bayesian classifier analysis, assesses membership probabilities for several nearby, young moving groups. According to BANYAN II, WISEA 1147–2040 has an 84.32% chance of belonging to TWA, under the assumption that it is young.

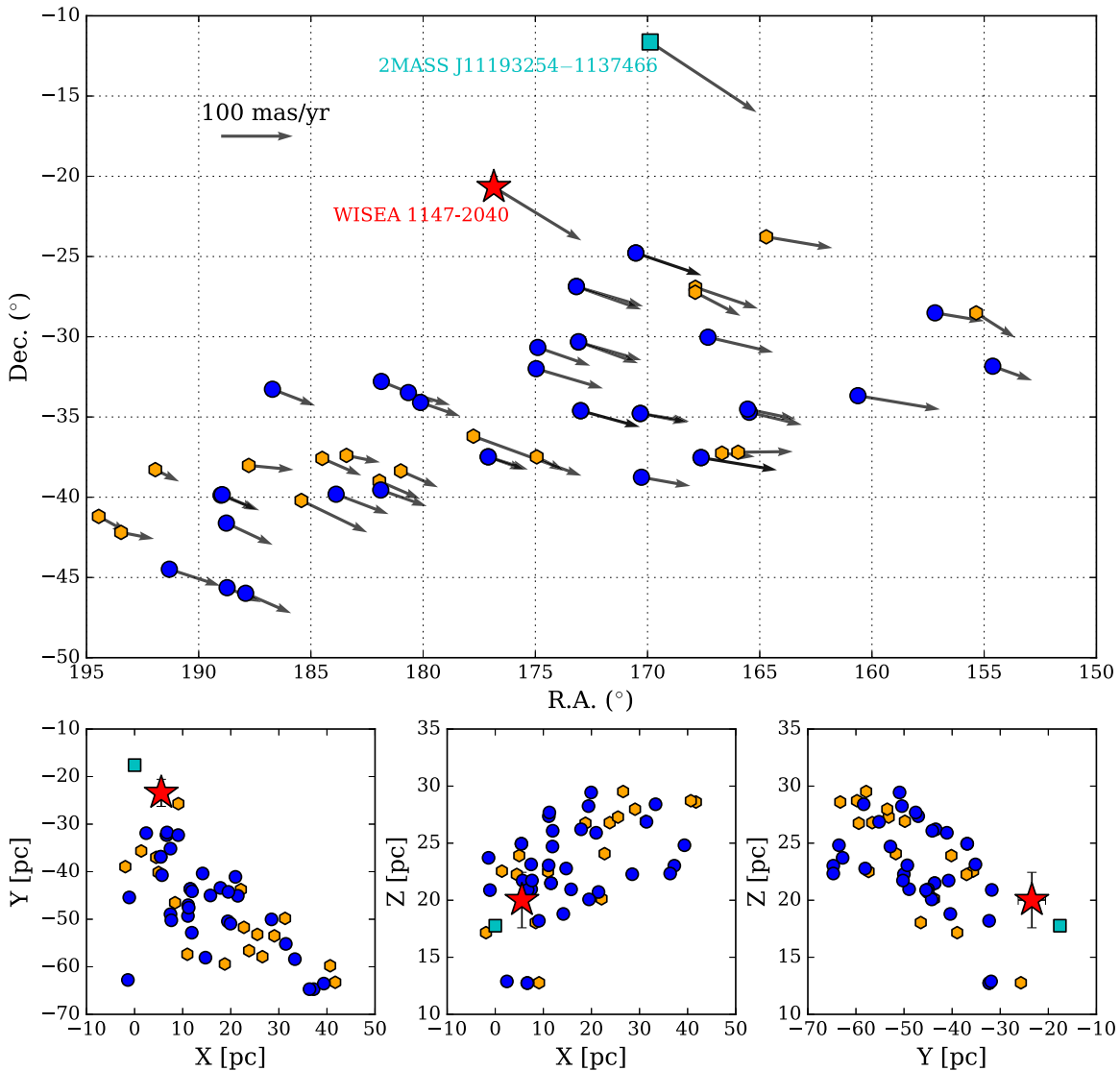
We can first evaluate the feasibility of TWA membership for WISEA 1147–2040 by inspecting its sky position relative to other TWA members, as this particular association is confined to a particular area of the sky (compared to most other young moving groups). The top panel in Figure 4 shows the position and proper motion vector of WISEA 1147–2040 along with confirmed TWA members from Schneider et al. (2012a, 2012b) and Murphy et al. (2015), as well as high-probability ( $>50\%$ ) candidate TWA members from Gagné et al. (2015) and the recently announced L7 candidate TWA member 2MASS J11193254-1137466 from Kellogg et al. (2015). The figure shows that WISEA 1147–2040 is in close proximity to the other members of TWA, and is thus viable as a TWA candidate member. BANYAN II also provides predicted radial velocity and distance values, assuming WISEA 1147–2040 is a TWA member, of  $9.61 \text{ km s}^{-1}$  and  $31.3 \pm 3.8 \text{ pc}$ , respectively. While a higher-resolution spectrum will be required to measure the radial velocity of WISEA 1147–2040, we can compare photometric and kinematic distance estimates to those predicated by BANYAN II to see if they are in agreement.

We estimate the distance to WISEA 1147–2040 in two ways. First, we estimate its distance photometrically. Note that because WISEA 1147–2040 has such red near-infrared colors, the absolute magnitude–spectral type relations for field brown dwarfs cannot be used for WISEA 1147–2040 with all available photometric bands. However, Faherty et al. (2013) show that the young, very red L5 $\gamma$  brown dwarf 2MASS J035523.37+113343.7 and the field L5 dwarf 2MASS

J1507476-162738 have very similar absolute flux values at the  $K$  band, while 2MASS J035523.37+113343.7 emits less flux than 2MASS J1507476-162738 at wavelengths shorter than  $K$ , and emits more flux at wavelengths redward of  $K$ . We investigated whether or not this trait applies to other young, very red L7 dwarfs with measured parallaxes, specifically PSO J318.5338–22.8603 (Liu et al. 2013) and WISEP J004701.06+680352.1 (Gizis et al. 2012). Using the absolute magnitude–spectral type relations of Dupuy & Liu (2012), we calculate a  $K_{\text{MKO}}$  photometric distance for PSO J318.5338–22.8603 and a spectral type of L7 of  $25.1 \pm 0.2 \text{ pc}$ , which agrees quite well with the measured parallax distance of  $24.6 \pm 1.4 \text{ pc}$  (Liu et al. 2013), especially when compared to the distances found using the  $J_{\text{MKO}}$  ( $39.5 \pm 0.7 \text{ pc}$ ) and  $W2$  ( $17.9 \pm 0.2 \text{ pc}$ ) band relations. Similarly, for WISEP J004701.06+680352.1 and a spectral type of L7, we calculate a  $K_S$  photometric distance of  $13.5 \pm 0.2 \text{ pc}$ , agreeing well with the measured distance of  $12.2 \text{ pc}$  (Gizis et al. 2015). Again, the  $K_S$  photometric distance is much more accurate than the  $J$  ( $18.5 \pm 0.6 \text{ pc}$ ) and  $W2$  ( $10.3 \pm 0.1 \text{ pc}$ ) band photometric distances. Therefore, we use the  $K_S$  absolute magnitude–spectral type relation of Dupuy & Liu (2012) to estimate a photometric distance to WISEA 1147–2040 of  $31.2 \pm 1.5 \text{ pc}$ .

We also estimate a kinematic distance to WISEA 1147–2040 following the “moving cluster” or “convergent point” method outlined in Mamajek (2005). This method works because the proper motions of comoving stars appear to converge to a single point on the celestial sphere from Earth’s frame of reference. Mamajek (2005) provides coordinates of the convergent point for TWA of ( $\alpha_{\text{cp}} = 103.2 \pm 1.5$ ,  $\delta_{\text{cp}} = -30.7 \pm 1.5$ ), while Ducourant et al. (2014) find ( $\alpha_{\text{cp}} = 102.4 \pm 1.4$ ,  $\delta_{\text{cp}} = -27.3 \pm 0.6$ ). For this analysis, we evaluate the kinematic distance to WISEA 1147–2040 using both sets of convergent point coordinates. We use mean  $UVW$  space velocities for TWA members of ( $U, V, W$ ) = ( $-11.12 \pm 0.90$ ,  $-18.88 \pm 1.56$ ,  $-5.63 \pm 2.78$ )  $\text{km s}^{-1}$  from Gagné et al. (2014). For WISEA 1147–2040, we find kinematic distances of  $32.6 \text{ pc}$  and  $32.0 \text{ pc}$  for the Mamajek (2005) and Ducourant et al. (2014) convergent points, respectively. Radial velocities can also be estimated from the convergent point method, for which we find values of  $9.2$  and  $8.7 \text{ km s}^{-1}$  for the Mamajek (2005) and Ducourant et al. (2014) convergent points, respectively. These radial velocity values agree very well with those estimated from BANYAN II.

Both the photometric distance from the  $K_S$  magnitude and the kinematic distance estimates of WISEA 1147–2040 agree remarkably well with each other, as well as the predicted distance from BANYAN II. If we use either the photometric or kinematic distance to WISEA 1147–2040 as an additional input into BANYAN II, we find TWA membership probabilities of  $\sim 96\%$ . Using the BANYAN II predicted distance, we compare the galactic  $XYZ$  coordinates of WISEA 1147–2040 with bona fide and high-probability candidate members of TWA in the bottom panel of Figure 4. Measured parallaxes for TWA members come from Weinberger et al. (2013) and Ducourant et al. (2014), where distances with the smallest uncertainties were chosen for objects found in both studies. We use the kinematic distances to TWA 3A, TWA 6, TWA 30A, and TWA 30B from Ducourant et al. (2014) and the kinematic distances to TWA 33 and TWA 34 from Schneider et al. (2012b). We use the quoted BANYAN distances to TWA 35 and TWA 36 from Murphy et al. (2015) and the BANYAN



**Figure 4.** Top: the position of WISEA 1147–2040 (red star) in the sky relative to known TWA members (blue circles), high-probability (>50%) candidates from Gagné et al. (2015; orange hexagons), and 2MASS J11193254–1137466 (light blue square) from Kellogg et al. (2015). Bottom: the XYZ positions of WISEA 1147–2040, known TWA members, high-probability (>50%) candidates from Gagné et al. (2015), and 2MASS J11193254–1137466. Symbols are the same as the upper panel.

distances to high probability candidates from Table 4 of Gagné et al. (2015b) come from Gagné et al. (2015a). We also include the L7 TWA candidate member 2MASS J11193254–1137466 from Kellogg et al. (2015), using their distance estimate of  $\sim 25$  pc. The figure shows that WISEA 1147–2040 has XYZ positions consistent with other TWA members. Based on WISEA 1147–2040’s young age, sky position, and the excellent agreement between its photometric distance estimate, kinematic distance estimate, and its predicted distance from BANYAN II, we conclude that WISEA 1147–2040 is a member of the TWA.

## 5. DISCUSSION

Based on WISEA 1147–2040’s  $T_{\text{eff}}$  and  $\log g$  estimates from its spectrum and an age of  $10 \pm 3$  Myr for TWA (Bell et al. 2015), we estimate a mass of  $6\text{--}13 M_{\text{Jup}}$  and  $9\text{--}11 M_{\text{Jup}}$  from the COND evolutionary models (Baraffe et al. 2003) and the Saumon & Marley (2008) ( $f_{\text{sed}} = 2$ ) models, respectively. Note that Liu et al. (2013) conclude that physical properties for

PSO J318.5338–22.8603 (and other young, red L dwarfs) derived from near-infrared spectra are unreliable, instead using a combination of near-infrared spectra, photometry, and a distance measurement to determine a bolometric luminosity. Mass, effective temperature, and surface gravity were then determined using the measured bolometric luminosity and moving group age.

While a parallax measurement for WISEA 1147–2040 is yet unavailable, we can use the distance estimate from BANYAN II ( $31.3 \pm 3.8$  pc), combined with our flux-calibrated near-infrared SpeX spectrum and WISE W1 and W2 magnitudes to estimate a preliminary bolometric luminosity. Any wavelengths not covered by the SpeX near-infrared spectrum or AllWISE photometry we fill in with the best-fitting model from Figure 2. We Monte Carlo the uncertainties for the flux-calibrated spectrum, AllWISE photometry, and distance estimate to determine the uncertainty of the calculated luminosity. We find  $\log(L_{\text{bol}}/L_{\odot}) = -4.42 \pm 0.11$ , which corresponds to  $T_{\text{eff}}$  values of  $\sim 1100\text{--}1200$  K at an age of  $10 \pm 3$  Myr (Saumon &

Marley 2008), much lower than the  $T_{\text{eff}}$  of WISEA 1147–2040 found via model fitting. This luminosity corresponds to a mass estimate of 5–6  $M_{\text{Jup}}$ . The differences between the parameters derived via fitting models to the near-infrared spectrum and those determined from the bolometric luminosity are similar to the differences seen for PSO J318.5338–22.860 in Liu et al. (2013).

Both the mass estimates from model fitting and from the bolometric luminosity make WISEA 1147–2040 the lowest-mass free-floating confirmed member of the TWA and one of the lowest-mass brown dwarfs in the solar neighborhood. In the TWA, only the planetary mass companion 2M1207b (Chauvin et al. 2004, 2005) has a lower mass. As such, WISEA 1147–2040 provides an exceptional laboratory for investigating the chemistry and cloud structure in young, planetary mass objects. A higher-resolution spectrum of WISEA 1147–2040 would provide both a radial velocity, helping to secure TWA membership, as well as additional gravity sensitive diagnostics (e.g., K I equivalent widths). A trigonometric parallax for WISEA 1147–2040 would also further confirm its membership in TWA.

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