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Young Accreting Compact Objects in M31: The Combined Power of *NuSTAR*, *Chandra*, and *Hubble*

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Abstract

We present 15 high-mass X-ray binary (HMXB) candidates in the disk of M31 for which we are able to infer compact object type, spectral type of the donor star, and age using multiwavelength observations from *NuSTAR*, *Chandra*, and the *Hubble Space Telescope*. The hard X-ray colors and luminosities from *NuSTAR* permit the tentative classification of accreting X-ray binary systems by compact object type, distinguishing black hole from neutron star systems. We find hard-state black holes, pulsars, and non-magnetized neutron stars associated with optical point-source counterparts with similar frequency. We also find nine non-magnetized neutron stars coincident with globular clusters and an equal number of pulsars with and without point-source optical counterparts. We perform spectral energy distribution (SED) fitting for the most likely optical counterparts to the HMXB candidates, finding seven likely high-mass stars and one possible red helium-burning star. The remaining seven HMXB optical counterparts have poor SED fits, so their companion stars remain unclassified. Using published star formation histories, we find that the majority of HMXB candidates—X-ray sources with UV-bright point-source optical counterpart candidates—are found in regions with star formation bursts less than 50 Myr ago, and three are associated with young stellar ages (<10 Myr). This is consistent with similar studies of HMXB populations in the Magellanic Clouds, M33, NGC 300, and NGC 2403.

Key words: galaxies: individual (M31) – pulsars: general – stars: black holes – stars: neutron – X-rays: binaries – X-rays: galaxies

1. Introduction

The production of extragalactic X-ray binaries (XRBs) is closely related to properties of the galaxies in which they form, such as the star formation rate (e.g., Ranalli et al. 2003; Gilfanov et al. 2004; Mineo et al. 2012), stellar mass (Lehmer et al. 2010), and metallicity (Basu-Zych et al. 2013; Brorby et al. 2016). Population studies of XRBs probe the production of these compact objects and their relationship to their host galaxy properties. However, the fundamental properties of XRBs, such as the compact object type and the physical properties of the donors, have remained difficult to determine given the limited information contained in the 0.5–10 keV energy range covered by soft X-ray telescopes such as *Chandra* and *XMM-Newton*. Broadening the observed energy range to include data from the near-IR (*Hubble Space Telescope (HST)*) through hard X-rays (*NuSTAR*) allows us to determine the compact object type and the physical properties of the donors, and to place constraints on the age of XRBs using star formation histories (SFHs) for their surrounding stellar populations.

Conducting a galaxy-wide study of high-mass X-ray binaries (HMXBs) in connection to their star-forming environments is challenging in the Milky Way because of the distance uncertainties, but there has been some successful work in this area (Grimm et al. 2002). Additionally, Bodaghee et al. (2012) used the spatial correlation between HMXBs and OB associations in the Milky Way to determine ages of the systems. It is expected that 5–10 Myr elapse between the formation of a high-mass star and the supernova that forms the compact object in HMXBs (Schaller et al. 1992; Linden et al. 2010). Thus, an HMXB cannot migrate far from its birthplace, allowing its spatial correlation with an OB association to be used to constrain its age. Bodaghee et al. (2012) determined the time from supernova through the HMXB phase (the "kinematic age") using the spatial correlation between HMXB candidates and OB associations. They found that most systems have kinematic ages of ∼4 Myr.

Detailed studies of XRBs in extragalactic star-forming environments have been made previously in the Small Magellanic Cloud (SMC), Large Magellanic Cloud (LMC), NGC 300, NGC 2403, and M33. In the SMC, B*e*/X-ray binaries are found in regions with star-forming bursts 25–60 Myr ago (Antoniou et al. 2010). In the LMC, HMXBs are found in areas with considerably more recent star formation, between 6 and 25 Myr ago (Antoniou & Zezas 2016). In NGC 300 and NGC 2403, HMXB candidates have been found in regions with surrounding stellar populations between 20 and 70 Myr old (Williams et al. 2013) with a peak

at 40–55 Myr, which agrees with ages in the SMC. In M33, a similar set of peaks is seen in the HMXB age distribution (Garofali et al. 2018). These ages suggest two potential formation channels: one that operates on the timescale of B-star evolution (∼50 Myr), and another that operates much more promptly.

To better connect the properties of the XRBs themselves to their parent populations, classifying the compact object in the system is critical. However, compact object characterization can be difficult because there are currently very few methods available. If a low-mass XRB has an observed Type-I X-ray burst, its compact object may be classified as a neutron star (e.g., Lewin et al. 1993). Black holes can be classified as such if the mass of the compact object can be confirmed using the orbital period and mass of the companion (Orosz & Bailyn 1997), but stellar companion orbits are not always available, especially for extragalactic XRBs.

The hard X-ray coverage of *NuSTAR* allows compact objects to be tentatively classified using their X-ray properties. With X-ray observations that cover the hard band (4–25 keV), compact objects can be classified as neutron stars or black holes based on a combination of their X-ray colors and luminosities (Wik et al. 2014; Yukita et al. 2016; A. Zezas et al. 2018, in preparation). This can be done because neutron stars always have hard emission associated with matter accreting onto the surface, while black holes do not and are dominated by the disk emission properties (Maccarone et al. 2016). Thus, an XRB's hard X-ray colors distinguish between neutron star and black hole systems. Techniques involving hard X-rays are more indirect but are critical to expand our tool kit for classifying X-ray sources as black holes or neutron stars.

Andromeda (M31), the nearest spiral galaxy to the Milky Way, is one of the best systems for studying X-ray binary populations in the context of their star-forming environments because of its proximity and the large number of multi-wavelength data sets available. Observations with the sensitivity to detect faint point sources extend from near-IR wavelengths up to hard X-rays ($E \lesssim 50$ keV; e.g., Vulic et al. 2014, 2016; Williams et al. 2014; Maccarone et al. 2016; Yukita et al. 2017).

There has recently been a major improvement in the X-ray coverage of Andromeda owing to two powerful and deep surveys by *NuSTAR* and *Chandra*, both taken in 2015. *NuSTAR* observed an ∼750 arcmin^2 area of M31 with a total exposure time of ∼1.4 Ms and average depth of ∼400 ks (D. Wik et al. 2018, in preparation). A *Chandra* Large Project survey covered ∼1800 arcmin^2 to a depth of 50 ks (ChandraPHAT; Williams et al. 2018).

We pair these X-ray observations with existing near-IR to UV observations from *Hubble* (PHAT; Dalcanton et al. 2012; Williams et al. 2014) to study hard X-ray emitting compact objects and their optical counterparts in the context of their star-forming environments. A total area of ∼570 arcmin^2 is covered by all three telescopes. This area comprises approximately 6% of the D_{25} area of M31.

The maturity of the PHAT project means that invaluable secondary data products are available to characterize the star-forming environments around X-ray sources. For example, Lewis et al. (2015) spatially mapped the recent SFH of M31 and Gregersen et al. (2015) mapped the metallicity distribution. Both properties allow the X-ray binary population to be placed in the context of its environment. The code called Bayesian extinction and stellar tool (BEAST) by Gordon et al. (2016) can fit the spectral energy distribution (SED) of individual stars in the disk of M31. The BEAST code provides a powerful tool for understanding the physical characteristics of the companion star in an XRB. Additionally, M31 allows us to study XRB populations in their environments without the uncertainties in the distance to each system that plague such studies in the Milky Way.

In this paper we use the multiwavelength coverage from *NuSTAR*, *Chandra*, and *HST* to investigate the HMXB population in the northern disk of M31. In Section 2 we describe the three data sets used in this study: *NuSTAR* observations, two sets of *Chandra* observations, and reduced *HST* photometry and imaging from the PHAT survey. We describe the methods used to match sources between the three data sets in Section 3. In Section 4 we describe our results: how *NuSTAR* sources were classified using their X-ray colors and luminosities, how we determined ages for HMXB candidates using spatially resolved SFHs, and the SED fitting used to determine spectral types for companion stars in HMXB candidates. In Section 5 we discuss our results in the context of previous studies, and in Section 6 we provide a brief summary of our results.

We assume a Galactic column density, $N_H = 7 \times 10^{20}$ cm^{-2}, and a photon index, $\Gamma = 1.7$ (Stiele et al. 2011), to convert count rates into absorbed energy flux. We assume a distance of 776 kpc to M31 (Dalcanton et al. 2012) for luminosity calculations.

2. Data

In this study we employ data from *NuSTAR*, *Chandra*, and *HST*. We now describe each data set in more detail below. For an overview of the area observed by each telescope, see Figure 1.

2.1. NuSTAR *Data*

NuSTAR source catalogs and source classifications come from Wik et al. (2018, in preparation). Observations were taken between February and October 2015 covering the area outlined in green in Figure 1 with an average exposure time of ∼400 ks, and they were reduced using the `nupipeline` software. Sources previously observed with *Chandra* were used for astrometric alignment. For detailed information on data reduction, source detection, and source classification, see Wik et al. (2018, in preparation), who present the entire *NuSTAR* M31 survey. The *NuSTAR* observations cover the nucleus and inner disk regions of M31 at an energy range of 4–25 keV using the 4–6 keV, 6–12 keV, and 12–25 keV energy bands. The completeness of the *NuSTAR* observations starts to fall off at a luminosity of $\sim 3 \times 10^{36}$ erg s^{-1} and reaches zero at $\sim 2 \times 10^{36}$ erg s^{-1}.

2.2. Chandra *Data*

Chandra data used in this study are comprised of two data sets: the ChandraPHAT data, a *Chandra* Large Project survey by B. Williams et al. (2018, submitted), and one additional *Chandra* field (obsid 18046, P.I. Hornschemeier), hereby referenced as Field A, that was reduced for this paper. The ChandraPHAT data set is comprised of seven *Chandra* pointings, each with a depth of about 50 ks. The Field A data consist of one *Chandra* pointing with a depth of 25 ks. The

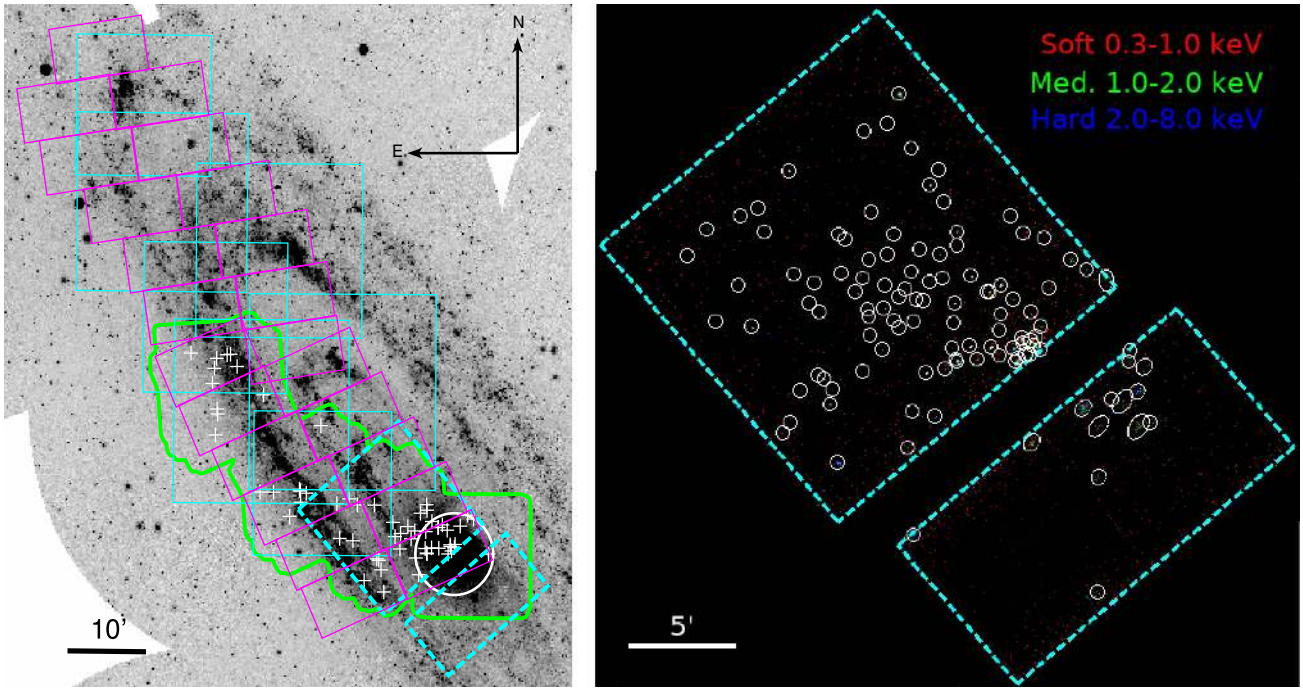

Figure 1. Near-UV image of M31 from the *Galaxy Evolution Explorer* (*GALEX*; left; Gil de Paz et al. 2009) and three-color X-ray image of *Chandra* Field A (right), see Section 2.2 for more information on Field A data reduction. Magenta regions outline the area observed by PHAT. Green outlines the *NuSTAR* observed region, and cyan outlines the area observed by *Chandra*, with solid lines indicating ChandraPHAT observations and dashed lines indicating Field A observations. In the UV image, the 64 sources observed by *NuSTAR* and *Chandra* that fall within the PHAT footprint are marked with white crosses, and a 5′ circle outlines the densest inner bulge region of M31. In the X-ray image, all sources detected by wavdetect within Field A are marked with white ellipses (including sources that do not match *NuSTAR* sources or are outside the PHAT footprint, and thus are not presented in our sample in Table 1).

ChandraPHAT observations were taken in 2015 October and the Field A observation was taken in 2016 August. The completeness in the ChandraPHAT field starts to drop at a luminosity of $\sim 3 \times 10^{35}$ erg s^{-1} and reaches zero at $\sim 5 \times 10^{34}$ erg s^{-1}. The completeness in Field A starts to drop at $\sim 7 \times 10^{35}$ erg s^{-1} and reaches zero at $\sim 1 \times 10^{35}$ erg s^{-1}.

Field A is centered at (R.A., decl.) = (00:43:29.30, +41:18:21.20) and was designed to overlap with *NuSTAR* Field A in the observations by Wik et al. (2018, in preparation). For a detailed description of the data reduction for *Chandra* sources within the ChandraPHAT footprint, see Williams et al. (2018, submitted). Sections 2.2.1 and 2.2.2 contain a detailed description of the reduction of the *Chandra* Field A data, which follows the methodology in Williams et al. (2018, submitted).

2.2.1. Source Detection

We generated the initial source list using the `wavdetect` tool of CIAO version 4.9 and CALDB version 4.7.7 (Fruscione et al. 2006). We produced the source image and exposure map using the CIAO command `fluximage` and created the PSF map using the CIAO tool `mkpsfmap` with the standard parameters, energy = 1.4967 and ecf = 0.393. We then ran `wavdetect` using the source image and PSF map to create a source list. The wavelet scales were set to 1.0, 2.0, 3.0, 8.0, and 16.0 pixels.

We ran ACIS extract (AE) version 2016sept22 (Broos et al. 2010) on the source list output from `wavdetect`. We followed Section 3.2 of the AE users guide to prepare the event files, aspect histogram file, aspect solution file, and mask file for AE source extraction. With the input source list from `wavdetect`, we extracted sources with energy limits of 0.35–8.0 keV. See Table 1 for the positions, off-axis angle, and net counts output by AE.

We iterated AE four times in order to obtain the most precise positions. The first AE run used the output positions from CIAO `wavdetect` as the initial positions. Each subsequent AE run used the data mean position output from the previous run as the initial positions; this was repeated until the input and output data mean positions converged (Figure 2).

2.2.2. Astrometric Alignment

After we obtained precise positions for the *Chandra* sources with ACIS extract, we aligned the Field A observations to the PHAT data set (Dalcanton et al. 2012). *NuSTAR* data were previously aligned to the ChandraPHAT data by Wik et al. (2018, in preparation).

We used eight bright globular clusters that were detected at optical wavelengths by *HST* and at X-ray wavelengths by *Chandra* for astrometric alignment. Clusters were identified by visual inspection of the PHAT imaging. First, we measured the centroids of the clusters using the `centroid_1dg` tool in the `photutils` (v0.4) Python package. To find the astrometric solution, we used the CIAO tool `wcs_match` that aligns the *Chandra* sources in a given image to the measured cluster positions from the PHAT images and outputs an astrometric solution. The parameters used in `wcs_match` were radius = 5, residlim = 0, residtype = 0, and residfac = 25 using the description in Vulic et al. (2016) as a guide. The CIAO tool `wcs_update` was used to updated the header of the *Chandra* images and update the R.A. and decl. of the measured source positions.

Positional errors were calculated using the net counts in the 0.35–8.0 keV band and the off-axis angle using the formula in

Table 1
All Sources Observed by *NuSTAR* and *Chandra* within the PHAT Footprint

NuSTAR ID	*Chandra* Catalog Name	*Chandra* R.A.	*Chandra* Decl.	*Chandra* R.A. err (″)	*Chandra* Decl. err (″)	Theta (′)	*Chandra* Flux (0.35–8.0 keV) ($\times 10^{-13}$ erg cm^{-2} s^{-1})	Stiele ID	Stiele Class.	PHAT Cpt.
19	004220.96+411520.3	10.587329	41.255773	2.0	2.0	13.0	$4.2^{+0.4}_{-0.4}$	…	…	n
24	004231.27+411937.5	10.630292	41.327141	2.0	2.0	10.0	$1.9^{+0.3}_{-0.2}$	923	GlC	c
26	004235.20+412005.0	10.646687	41.334782	1.0	0.9	9.5	$3.4^{+0.3}_{-0.3}$	952	⟨hard⟩	n
27	004240.31+411845.3	10.667979	41.312682	2.0	1.0	8.6	$0.90^{+0.15}_{-0.14}$	972	⟨hard⟩	p
41	004243.81+411631.0	10.682538	41.275371	0.7	0.4	8.3	$13.0^{+0.7}_{-0.6}$	1005	⟨XRB⟩	n
43	004244.27+411607.6	10.684498	41.268868	0.7	0.4	8.4	$19.0^{+0.9}_{-0.9}$	1010	⟨XRB⟩	n
44	004246.19+411543.2	10.692552	41.262086	2.0	2.0	8.2	$0.76^{+0.25}_{-0.21}$	…	…	n
45	004246.97+411615.3	10.695737	41.271019	0.6	0.4	7.9	$13.0^{+0.6}_{-0.6}$	1023	⟨AGN⟩	n
46	004247.18+411628.0	10.696608	41.274542	0.6	0.3	7.8	$26.0^{+0.8}_{-0.8}$	1024	⟨XRB⟩	n
47	004248.56+411520.8	10.702332	41.255889	0.6	0.3	8.0	$51.0^{+1.0}_{-1.0}$	1036	⟨XRB⟩[a]	n
54	004249.22+411815.5	10.70509	41.304395	0.8	0.5	6.9	$1.5^{+0.2}_{-0.1}$	1041	⟨hard⟩	p
55	004252.53+411854.0	10.718888	41.315082	0.5	0.2	6.3	$19.0^{+0.5}_{-0.5}$	1060	⟨XRB⟩	n
65	004254.93+411602.8	10.728903	41.267544	0.5	0.2	6.6	$16.0^{+0.5}_{-0.5}$	1075	⟨XRB⟩[a]	n
57	004255.19+411835.4	10.729964	41.309922	0.7	0.5	5.8	$0.83^{+0.12}_{-0.10}$	1078	⟨hard⟩	n
57	004255.60+411834.5	10.731708	41.309676	0.7	0.4	5.7	$0.91^{+0.12}_{-0.11}$	1078	⟨hard⟩	c
59	004259.66+411918.9	10.748579	41.322021	0.4	0.2	4.9	$11.0^{+0.4}_{-0.4}$	1102	GlC	c
68	004259.88+411605.3	10.749525	41.268225	0.5	0.2	5.8	$10.0^{+0.4}_{-0.4}$	1103	GlC	c
70	004302.94+411522.2	10.762284	41.256263	0.5	0.2	5.8	$6.4^{+0.3}_{-0.3}$	1116	GlC	c
77	004303.03+412041.6	10.762621	41.344983	0.6	0.4	4.5	$0.50^{+0.09}_{-0.07}$	1115	⟨hard⟩	n
70	004303.23+411527.3	10.763472	41.257677	0.5	0.2	5.7	$13.0^{+0.4}_{-0.4}$	1116	GlC	n
78	004303.29+412121.5	10.763713	41.35607	0.5	0.2	4.7	$2.3^{+0.2}_{-0.2}$	1118	GlC	c
60	004303.87+411804.5	10.766124	41.301336	0.4	0.2	4.3	$6.0^{+0.3}_{-0.3}$	1122	GlC	c
71	004304.25+411600.7	10.767706	41.266967	0.7	0.5	5.2	$0.59^{+0.10}_{-0.09}$	1124	⟨GlC⟩	n
79	004307.51+412019.4	10.781315	41.338806	0.5	0.3	3.6	$0.58^{+0.09}_{-0.08}$	1137	⟨GlC⟩	c
80	004308.62+411248.0	10.785948	41.213448	0.8	0.5	7.2	$3.2^{+0.3}_{-0.3}$	1146	XRB	p
81	004310.62+411451.0	10.794248	41.247599	0.4	0.2	5.2	$23.0^{+0.6}_{-0.6}$	1157	GlC	c
82	004311.37+411809.3	10.797389	41.302675	0.5	0.2	2.9	$0.78^{+0.10}_{-0.09}$	1160	⟨hard⟩	n
85	004313.88+411711.5	10.807835	41.286615	0.7	0.5	3.0	$0.13^{+0.05}_{-0.04}$	…	…	n
86	004316.10+411841.2	10.817115	41.311543	0.4	0.2	1.9	$0.35^{+0.06}_{-0.07}$	1180	⟨XRB⟩	p
88	004321.07+411750.2	10.837815	41.297389	0.4	0.1	1.7	$1.1^{+0.1}_{-0.1}$	1203	⟨hard⟩	p
87	004321.48+411556.5	10.839501	41.265805	0.7	0.4	3.4	$0.21^{+0.05}_{-0.06}$	…	…	p
89	004324.84+411726.9	10.853509	41.290917	0.4	0.2	1.8	$0.47^{+0.08}_{-0.07}$	1216	⟨hard⟩	n
90	004326.33+411911.4	10.859718	41.31994	0.4	0.1	0.11	$0.38^{+0.09}_{-0.07}$	1224	⟨AGN⟩	g
91	004332.38+411040.9	10.884951	41.178136	0.7	0.4	8.7	$15.0^{+0.5}_{-0.5}$	1253	⟨hard⟩[a]	n
92	004334.33+411323.1	10.893064	41.223187	0.5	0.3	6.1	$5.4^{+0.3}_{-0.3}$	1261	⟨hard⟩	n
93	004335.91+411433.0	10.899635	41.2426	0.8	0.6	5.1	$0.41^{+0.09}_{-0.07}$	1262		p
94	004337.28+411443.1	10.905322	41.245424	0.4	0.2	5.0	$16.0^{+0.4}_{-0.4}$	1267	GlC	c
95	004339.06+412116.7	10.912737	41.354885	0.7	0.7	7.2	$0.73^{+0.07}_{-0.06}$	…	…	p
96	004345.83+411203.7	10.940976	41.201128	2.0	2.0	8.1	$0.39^{+0.11}_{-0.09}$	1298	⟨hard⟩	n
99	004350.76+412117.4	10.961516	41.355033	0.4	0.4	5.1	$0.65^{+0.06}_{-0.05}$	1319	⟨hard⟩	p
97	004353.65+411654.6	10.973526	41.282044	0.4	0.4	7.6	$6.6^{+0.2}_{-0.2}$	1327	⟨GlC⟩	n
98	004356.43+412202.3	10.985126	41.367503	0.4	0.4	3.8	$0.45^{+0.05}_{-0.05}$	1340	GlC	c
105	004402.72+411711.3	11.011322	41.28666	1.0	1.0	6.5	$0.18^{+0.04}_{-0.03}$	…	…	n
101	004404.75+412126.5	11.019799	41.35756	0.3	0.3	2.7	$0.75^{+0.06}_{-0.06}$	1373	⟨AGN⟩	p
102	004416.02+413057.3	11.066667	41.516147	0.4	0.4	4.3	$1.7^{+0.1}_{-0.1}$	1420	XRB	n
103	004425.73+412241.8	11.107221	41.378442	0.3	0.3	2.0	$0.22^{+0.03}_{-0.03}$	1450	⟨hard⟩	p
104	004429.57+412135.1	11.123203	41.359913	0.3	0.3	3.1	$18.0^{+0.3}_{-0.3}$	1463	GlC	c
105	004429.73+412257.4	11.123878	41.382771	0.5	0.5	2.6	$0.09^{+0.02}_{-0.02}$	…	…	n
109	004430.16+412301.1	11.125694	41.383802	0.6	0.6	2.7	$0.04^{+0.02}_{-0.01}$	…	…	g
105	004430.45+412310.1	11.126901	41.3863	0.4	0.4	2.8	$0.21^{+0.03}_{-0.03}$	1468	⟨hard⟩	n
106	004437.08+411951.1	11.154504	41.331024	0.5	0.5	5.3	$0.42^{+0.05}_{-0.05}$	1488	⟨hard⟩	g
100	004448.13+412247.4	11.200545	41.379973	0.7	0.7	6.1	$0.27^{+0.04}_{-0.04}$	1525	⟨hard⟩	p
110	004455.53+413440.3	11.231167	41.57808	0.3	0.3	2.7	$0.45^{+0.05}_{-0.04}$	1547	⟨AGN⟩	n
108	004457.39+412247.9	11.239115	41.380094	0.7	0.7	7.8	$1.0^{+0.1}_{-0.1}$	1553	⟨XRB⟩	n
111	004513.82+413806.4	11.307524	41.635323	0.8	0.8	6.1	$0.24^{+0.04}_{-0.04}$	1598	⟨hard⟩	g
112	004518.39+413936.0	11.326586	41.66018	0.5	0.5	4.4	$0.28^{+0.04}_{-0.04}$	1611	⟨hard⟩	p

Table 1
(Continued)

NuSTAR ID	*Chandra* Catalog Name	*Chandra* R.A.	*Chandra* Decl.	*Chandra* R.A. err (″)	*Chandra* Decl. err (″)	Theta (′)	*Chandra* Flux (0.35–8.0 keV) ($\times 10^{-13}$ erg cm^{-2} s^{-1})	Stiele ID	Stiele Class.	PHAT Cpt.
113	004520.74+413932.1	11.336316	41.659109	0.6	0.6	4.3	$0.14^{+0.03}_{-0.03}$	…	…	g
117	004526.86+413216.8	11.361729	41.538161	0.6	0.6	5.2	$0.27^{+0.04}_{-0.03}$	1631	⟨AGN⟩	n
118	004527.34+413253.5	11.363743	41.548363	0.4	0.4	5.4	$2.2^{+0.1}_{-0.1}$	1634	⟨hard⟩	g
114	004527.89+413904.9	11.366179	41.651539	0.4	0.4	4.3	$0.43^{+0.05}_{-0.04}$	1635	⟨hard⟩	p
119	004528.29+412943.4	11.367681	41.495538	0.4	0.4	6.0	$1.9^{+0.1}_{-0.1}$	1636	⟨hard⟩	p
115	004529.35+413751.6	11.37223	41.631176	1.0	1.0	5.6	$0.06^{+0.02}_{-0.02}$	…	…	g
116	004530.65+413559.8	11.377557	41.600135	0.9	0.9	7.1	$0.38^{+0.06}_{-0.05}$	1643	⟨hard⟩	g
120	004545.57+413941.5	11.439867	41.661701	0.3	0.3	4.7	$71.0^{+0.6}_{-0.6}$	1692	GIC	c

Notes. List of all *NuSTAR–Chandra* sources within the PHAT footprint. Optical counterparts to X-ray sources are listed in the *PHAT Cpt.* column: g = galaxy, n = no optical counterpart, c = cluster, and p = point source. See Section 3.2 for further discussion of optical counterpart determination. Stiele ID and classifications from Stiele et al. (2011).
[a] Sources have updated Stiele classifications from Stiele & Kong (2018).

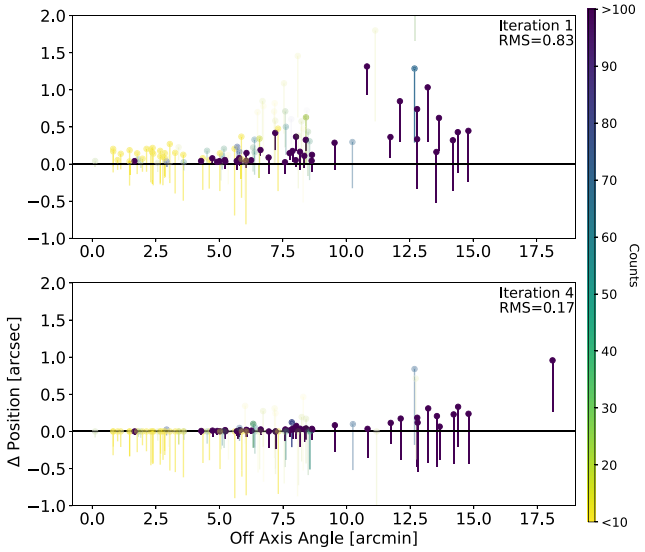

Figure 2. Results of ACIS extract iteration to improve *Chandra* source positions in Field A observations (see Section 2.2.1). The top panel shows the change in source position between the input source list from wavdetect and the output AE positions. The bottom panel shows the change between the input and output source positions for the fourth iteration of AE. Note that most of the sources with an off-axis angle >10′ have more than 100 counts. These belong to the nucleus of M31, which is far off axis in the Field A observations and has a high source density. The source with an off-axis angle of ∼17.′5 in the bottom panel is not shown in the top panel because it has a Δ position >2″ between the input position from wavdetect and the output position from AE.

Hong et al. (2005), listed as Equation (5). Instead of using the 0.″25 baseline error in that equation, we added the residuals from astrometric alignment to PHAT, which were 0.″29 in R.A. and 0.″03 in decl.

We merged the Field A catalog prepared for this paper with the ChandraPHAT catalog from Williams et al. (2018, submitted) for source matching with the *NuSTAR* source catalog by Wik et al. (2018, in preparation) and PHAT, detailed in Section 3.

2.3. *HST Data: Panchromatic* Hubble *Andromeda Treasury*

The *HST* data come from the published Panchromatic *Hubble* Andromeda Treasury (PHAT) data set (Dalcanton et al. 2012). The PHAT survey imaged roughly a third of the disk of M31 in six *HST* filters ranging from near-IR to UV wavelengths: F160W, F110W, F814W, F475W, F336W, and F275W (central $\lambda = 1.150\,\mu$, $1.545\,\mu$, 8353 Å, 4750 Å, 3375 Å, 2750 Å). We use published photometry catalogs by Williams et al. (2014) for optical counterpart analysis. PHAT observations were taken in 2010 and 2011. The PHAT data have a limiting F475W magnitude of ∼28 in the outer disk and ∼25 in the more crowded inner disk region.

3. Source Matching between Data Sets

We first matched *NuSTAR* sources to *Chandra* to find precise positions. We then used the *Chandra* positions to identify optical counterparts in the PHAT data.

3.1. *Source Matching between* NuSTAR *and* Chandra

We identified 60 *NuSTAR* sources with positions inside the PHAT observed area that positionally match *Chandra* sources. These 60 sources have 64 associated *Chandra*-detected X-ray sources. We cross-matched *NuSTAR* and *Chandra* sources within 10″ so we could use the more precise *Chandra* positions to identify optical counterparts.

We chose a 10″ match radius to account for the 9″ full width at half-maximum of the *NuSTAR* point-spread function (PSF) and the ∼0.″5 mean *Chandra* positional errors for sources in our sample. We measured the *Chandra* exposure time at the position of each detected *NuSTAR* source to confirm that if a source was observed by both telescopes, there was a match.

In order to quantify the confidence level of these matches, we investigated the false-match probability between *NuSTAR* and *Chandra*. To do this, we adjusted the *NuSTAR* source positions for the full 121-source *NuSTAR* M31 catalog (Wik et al. 2018, in preparation) by 10″ in both R.A. and decl. We performed this adjustment four times, using all permutations of adding and subtracting 10″ from the R.A. and decl. of *NuSTAR* sources. We re-matched the *NuSTAR* and *Chandra* source positions each time to see how many *Chandra* sources matched the adjusted *NuSTAR* source positions. We found an average of five matches between the adjusted *NuSTAR* source positions and the *Chandra* source positions. Out of 121 *NuSTAR* sources, this equals a false-match probability of 4.1%. This means that two to three of the *NuSTAR* sources in our sample could have false matches to *Chandra* sources.

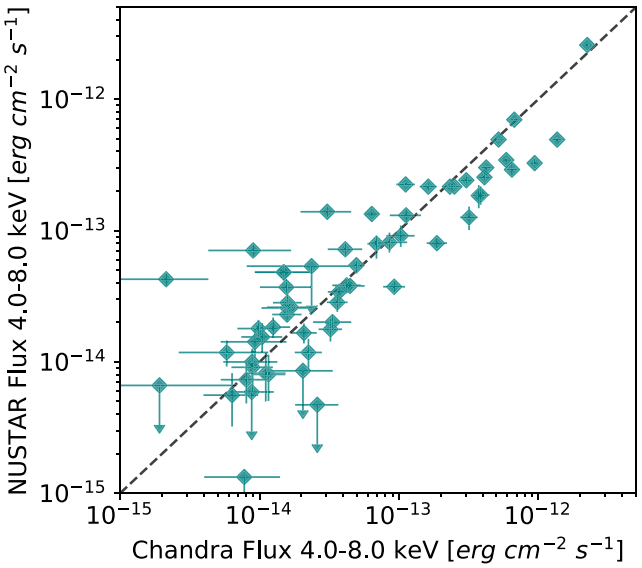

Figure 3. Comparison of *NuSTAR* and *Chandra* measured fluxes for 60 hard X-ray sources observed by both telescopes. Sources were matched positionally to within 10″.

We note that while our sample covers an area ∼75% of the *NuSTAR* total observed area, it only contains ∼50% of the *NuSTAR* sources in the full 121-source catalog. This is because our sample only contains sources observed with *NuSTAR*, *Chandra*, and *HST*, which excludes part of the bulge of M31, in an area with high *NuSTAR* source density.

Five *NuSTAR* sources lie within the *Chandra* Field A footprint that were not detected by *Chandra*. Given that the *Chandra* Field A and *NuSTAR* observations were taken a year apart, we believe this discrepancy is due to variability.

There are three *NuSTAR* sources whose positions are compatible with multiple *Chandra* sources. This is not surprising as *NuSTAR* can blend *Chandra* sources together because of its large PSF. The PSF of *NuSTAR* has a core with a full width at half-maximum of 18″ and a half-power diameter of 58″ (Harrison et al. 2013). This is quite large compared to the *Chandra* PSF, which is ∼0.″5 on-axis to ∼10″ at the edge of the field (Williams et al. 2004).

NuSTAR source 105 matched to three *Chandra* sources, and *NuSTAR* sources 70 and 57 matched to two *Chandra* sources. In these instances, we kept all *Chandra*-detected sources in our total list of 64 X-ray sources. When comparing *NuSTAR* classifications with optical counterpart types, we only used the optical counterpart associated with the *Chandra* source with the largest number of counts in the 0.35–8.0 keV band.

To further confirm associations between the *NuSTAR* and *Chandra* sources, we compared the flux of each match in the 4–8 keV energy range, as shown in Figure 3. We converted from count rates into fluxes for each telescope using the NASA High Energy Astrophysics Science Research Archive Center's (HEASARC) web-based Portable, Interactive Multi-Mission Simulator (WebPIMMS) tool.[12]

3.2. Identifying Optical Counterparts in PHAT

We determined optical counterparts to X-ray sources using a combination of positional matching, UV magnitude cuts, and visual inspection. We initially determined optical counterpart candidates by looking at both optical and UV images of the PHAT data within the 1σ *Chandra* positional error circles of a hard X-ray source. This method allowed for initial detection of likely counterparts such as background galaxies and globular clusters (e.g., Galleti et al. 2003). If a source had a clear point source in the UV F336W image within the 1σ *Chandra* positional errors, it was noted as a point-source counterpart candidate and its PHAT photometry was retrieved and is listed in Table 2.

We investigated the false-match probability for the PHAT counterparts. The PHAT survey area is divided into 23 "bricks" (see Dalcanton et al. (2012) for a description of brick boundaries). We calculated the source density of O/B stars in the PHAT survey in the bricks (9 bricks total) covered by *Chandra* and *NuSTAR* observations. We used only stars with good data in at least three of the *HST* photometric bands used by PHAT and an F336W magnitude lower than 23. Dividing the number of O/B stars by the total area of the 9 bricks gives a density of O/B stars per area. We then multiply this source density by the area of the average *Chandra* 1σ error circle to determine the probability of finding an O/B star within the 1σ error circle of an X-ray source. We find a false-match probability of about 2%. Accounting for the 64 X-ray sources in our combined *NuSTAR–Chandra* sample within the PHAT footprint, we expect one to two false matches.

We used finding charts and CMDs to identify optical counterparts. Figure 4 shows a representative figure for source 004335.91+411433.4, an X-ray source with a point-source optical counterpart. The optical counterpart is marked in the UV image (lower left) with a cyan circle and is visible as a very bright star in the 1σ *Chandra* positional errors of the optical finder (lower right). The counterpart is also plotted on two color–magnitude diagrams (CMDs) in the top row of the figure as a cyan star. It falls along the massive end of the main sequence in both CMDs. Note that there are far fewer stars in the upper left CMD because there are not as many stars in the PHAT survey that have well-measured UV (F336W) magnitudes. The very populated region of the upper right CMD is the red giant branch, which is too faint in the UV to be detected in the PHAT data, so that this feature is not as prominent in the UV CMD.

We looked for optical counterparts for the *NuSTAR–Chandra* sources, 64 of which are within the area of M31 observed by PHAT. We determined the following optical counterparts: 15 point sources, 13 globular clusters, and 8 background galaxies. The remaining 28 *NuSTAR–Chandra* sources do not have clear optical counterparts. Optical counterparts for all sources are listed in the last column of Table 1.

We expect to find seven to eight background AGN in our *NuSTAR* sample. Wik et al. (2018, in preparation) identified *NuSTAR* sources with luminosities greater than $\sim 2 \times 10^{36}$ erg s^{-1}, and used the published log(N)-log(S) relationship from Harrison et al. (2016) to calculate the expected contamination of background AGN. We scale this relation to the area of the *NuSTAR* field also covered by *Chandra* and *HST*. We identify eight background galaxies using the PHAT imaging (listed in Table 1), which is consistent with this prediction, suggesting that all AGN with *NuSTAR* detections were visible in the optical PHAT data.

[12] https://heasarc.gsfc.nasa.gov/cgi-bin/Tools/w3pimms/w3pimms.pl

Table 2
PHAT Photometry for HMXB Optical Counterpart Candidates

Catalog Name	*Chandra* R.A.	*Chandra* Decl.	PHAT R.A.	PHAT Decl.	F275W	F336W	F475W	F814W	F110W	F160W
004240.31+411845.3	10.667979	41.312682	10.667543	41.312546	22.03 ± 0.04	21.31 ± 0.01	20.940 ± 0.003	19.799 ± 0.003	19.263±0.003	18.512 ± 0.003
004249.22+411815.5	10.70509	41.304395	10.705035	41.304478	⋯	25.2 ± 0.2	22.119 ± 0.006	19.314 ± 0.002	18.552 ± 0.002	17.586 ± 0.002
004308.62+411248.0	10.785948	41.213448	10.785837	41.213419	23.19 ± 0.07	22.97 ± 0.03	23.41 ± 0.01	23.09 ± 0.02	23.26 ± 0.04	22.63 ± 0.05
004316.10+411841.2	10.817115	41.311543	10.817102	41.31151	24.1 ± 0.1	23.15 ± 0.04	24.43 ± 0.02	22.15 ± 0.01	21.43 ± 0.01	19.906 ± 0.007
004321.07+411750.2	10.837815	41.297389	10.837831	41.297466	24.9 ± 0.2	24.04 ± 0.06	24.43 ± 0.02	24.48 ± 0.06	25.0 ± 0.2	⋯
004321.48+411556.5	10.839501	41.265805	10.839394	41.265853	23.8 ± 0.1	23.20 ± 0.05	23.79 ± 0.01	23.19 ± 0.02	22.90 ± 0.03	21.72 ± 0.02
004335.91+411433.0	10.899635	41.2426	10.899774	41.242598	22.62 ± 0.05	22.80 ± 0.03	24.04 ± 0.01	23.53 ± 0.02	22.97 ± 0.04	23.29 ± 0.09
004339.06+412116.7	10.912737	41.354885	10.912961	41.354867	23.8 ± 0.1	23.23 ± 0.04	23.87 ± 0.01	23.63 ± 0.03	23.60 ± 0.04	23.27 ± 0.06
004350.76+412117.4	10.961516	41.355033	10.961508	41.355045	22.78 ± 0.06	21.16 ± 0.01	21.694 ± 0.004	19.843 ± 0.002	19.129 ± 0.002	18.252 ± 0.002
004404.75+412126.5	11.019799	41.35756	11.019758	41.357577	23.51 ± 0.09	22.36 ± 0.02	22.150 ± 0.005	19.938 ± 0.003	18.735 ± 0.002	17.790 ± 0.002
004425.73+412241.8	11.107221	41.378442	11.107175	41.378477	26.1 ± 0.6	24.9 ± 0.1	24.65 ± 0.02	22.62 ± 0.01	21.880 ± 0.009	21.45 ± 0.01
004448.13+412247.4	11.200545	41.379973	11.200584	41.380057	23.48 ± 0.08	23.39 ± 0.04	24.59 ± 0.02	22.95 ± 0.01	21.851 ± 0.009	20.448 ± 0.006
004518.39+413936.0	11.326586	41.66018	11.326624	41.660177	25.6 ± 0.4	24.8 ± 0.1	25.28 ± 0.03	24.70 ± 0.04	24.10 ± 0.05	23.27 ± 0.05
004527.89+413904.9	11.366179	41.651539	11.36618	41.651542	24.3 ± 0.2	22.81 ± 0.03	23.405 ± 0.009	20.870 ± 0.004	19.971 ± 0.003	18.819 ± 0.002
004528.29+412943.4	11.367681	41.495538	11.367687	41.495535	19.76 ± 0.01	19.202 ± 0.005	20.231 ± 0.002	19.014 ± 0.002	18.526 ± 0.001	17.754 ± 0.001

Note. PHAT photometry for all point-source optical counterparts to *NuSTAR* hard X-ray sources. Sources are identified by their *Chandra* catalog name, which corresponds to the *Chandra Catalog Name* column in Table 1. Ellipses indicate that the source was not detected in that filter.

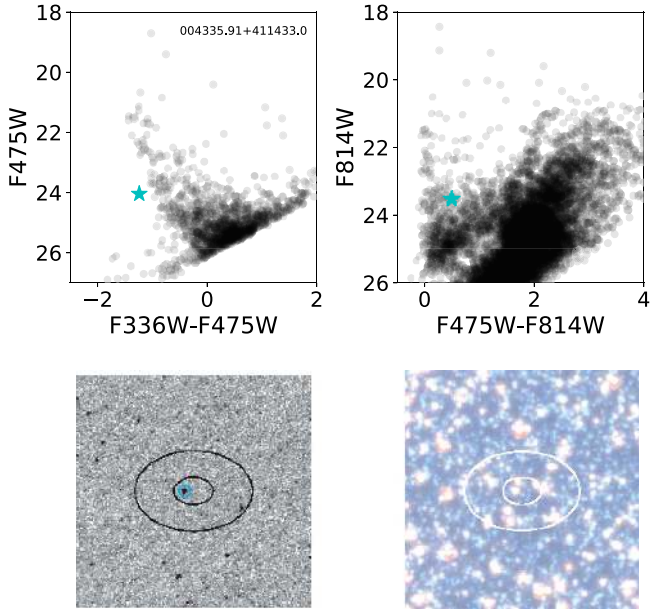

Figure 4. Color–magnitude diagrams (top) and finding charts (bottom) from the PHAT data set for the 10″ × 10″ region surrounding source 004335.91 +411433.4. The top two panels show the UV and optical CMDs. The bottom two panels show a UV image in the F336W filter and an RGB optical image with the F160W filter as red, F814W as green, and F336W as blue. The plotted ellipses represent the 1 and 3σ *Chandra* positional errors. The optical counterpart is identified in the UV finder (lower left) with a cyan circle and on the CMDs with a cyan star. The black points in the background represent other stars in the PHAT photometry catalog within 5″ of the X-ray source position.

4. Results

4.1. NuSTAR *Source Classification*

Wik et al. (2018, in preparation) classified the hard X-ray sources in this sample by comparing their X-ray colors and luminosities to those of Galactic XRBs with known compact object types. This method is presented in A. Zezas et al. (2018, in preparation) and has previously been used to classify sources in NGC 253 (Wik et al. 2014).

Black hole XRBs are known to exhibit different accretion states that are manifested by their different broad-band X-ray spectra (especially above 10 keV) and power-spectra (e.g., Done et al. 2004; Remillard & McClintock 2006). The main differences between these spectral states are identified at energies above 8.0 keV, i.e., energies that can be probed with *NuSTAR*.

In order to develop a diagnostic tool that can be used to characterize *NuSTAR* observations of extragalactic XRBs, Zezas et al. (2018, in preparation) used the extensive library of black hole spectra of Sobolewska et al. (2009). This library includes includes a set of 1772 Rossi-XTE—PCA observations of six Galactic black hole X-ray binaries. These observations were performed during different accretion states, and in some cases, they cover the complete evolution of a system during an outburst. Each spectrum was modeled with a Comptonized disk blackbody model (Sobolewska et al. 2009). The state characterization was based on the spectral shape (see Sobolewska et al. 2009, for more details).

Based on this model and the *NuSTAR* response files, A. Zezas et al. (2018, in preparation) simulated *NuSTAR* observations and calculated the expected count rates in

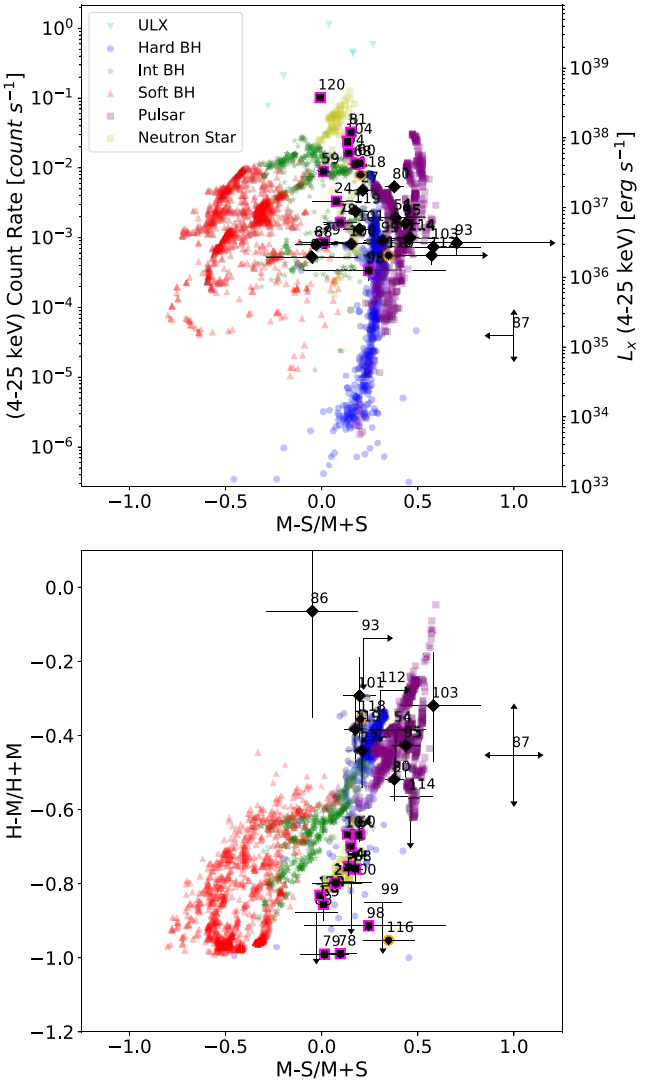

Figure 5. Hardness-intensity diagram and hardness-ratio plots used to classify *NuSTAR* sources. These plots show sources classified as HMXB candidates as black diamonds. Circles outlined in orange are the two background galaxies with *NuSTAR* classifications that were identified using *HST* imaging, described further in Section 5.1. Squares outlined in magenta are globular clusters, which occupy the non-magnetized neutron star region of both diagrams (see Section 5.5 for a further discussion of X-ray sources in clusters). The bands are defined as follows: soft (S = 4–6 keV), medium (M = 6–12 keV), and hard (H = 12–25 keV). Sources are labeled by *NuSTAR* ID. Background colored sources represent modeled evolutionary tracks of Galactic X-ray binaries with known compact object types, adjusted for the distance of M31 (see A. Zezas et al., 2018 in preparation for information on the Galactic XRB diagnostic regions). *NuSTAR* data and source classifications for M31 sources from Wik et al. (2018, in preparation).

different bands. Extensive tests showed that hardness ratios involving the 4.0–6.0 keV (soft), 6.0–12.0 keV (medium), 12.0–25.0 keV (hard), and 4.0–25.0 keV (full) bands give the optimal separation of spectral states, while maximizing the number of counts in each band. Luminosities and count rates of Galactic XRBs were scaled to the distance of M31 for comparison with XRBs in our sample.

We note that the highest energy of the *NuSTAR* data (25 keV) is well within the range of the RXTE-PCA spectra, ensuring high-quality input spectral models. Figure 5 shows the locus of the different black hole accretion states on the intensity-hardness ratio and hardness-ratio diagrams (red,

Table 3
Output Best-fit Parameters for Stellar Optical Counterparts from the BEAST SED Fitting Code

Catalog Name	$\log(L)$ ($L_\odot$)	$\log(g)$ (cm s^{-2})	$\log(T)$ (K)	A_V	Mass ($M_\odot$)	χ^2	Best-fit Spectral Type
004240.31+411845.6	$4.5^{+0.1}_{-0.1}$	$2.2^{+0.1}_{-0.1}$	$4.03^{+0.02}_{-0.02}$	$2.0^{+0.1}_{-0.1}$	13^{+1}_{-1}	49	χ^2 too high to trust fit
004249.22+411815.8	$3.8^{+0.1}_{-0.1}$	$0.7^{+0.2}_{-0.3}$	$3.61^{+0.02}_{-0.03}$	$0.7^{+0.3}_{-0.4}$	5^{+2}_{-3}	11	possible He burning star
004308.63+411248.4	$3.1^{+0.5}_{-0.3}$	$3.6^{+0.2}_{-0.1}$	$4.2^{+0.1}_{-0.1}$	$0.7^{+0.5}_{-0.3}$	5^{+2}_{-1}	1	B
004316.11+411841.5	$5.7^{+0.3}_{-1.1}$	$3.8^{+0.3}_{-0.8}$	$4.6^{+0.1}_{-0.4}$	$4.3^{+0.2}_{-0.3}$	45^{+34}_{-35}	50	high χ^2, flat SED
004321.08+411750.6	$3.0^{+0.5}_{-0.4}$	$4.1^{+0.2}_{-0.3}$	$4.2^{+0.1}_{-0.1}$	$1.2^{+0.4}_{-0.5}$	5^{+2}_{-1}	6	B
004321.48+411556.9	$3.6^{+0.5}_{-0.4}$	$4.1^{+0.3}_{-0.2}$	$4.4^{+0.1}_{-0.1}$	$1.2^{+0.2}_{-0.2}$	7^{+4}_{-3}	1	B
004335.91+411433.4	$4.0^{+0.3}_{-0.6}$	$4.2^{+0.2}_{-0.3}$	$4.5^{+0.1}_{-0.1}$	$1.7^{+0.2}_{-0.2}$	10^{+3}_{-4}	5	B
004339.06+412117.6	$3.4^{+0.4}_{-0.4}$	$4.1^{+0.2}_{-0.2}$	$4.3^{+0.1}_{-0.1}$	$1.0^{+0.2}_{-0.2}$	6^{+3}_{-2}	5	B
004350.76+412118.1	$6.5^{+0.1}_{-0.1}$	$3.8^{+0.1}_{-0.1}$	$4.72^{+0.02}_{-0.02}$	$3.2^{+0.1}_{-0.1}$	106^{+8}_{-7}	1756	χ^2 too high to trust fit
004404.75+412127.2	$5.2^{+0.1}_{-0.1}$	$1.9^{+0.1}_{-0.1}$	$4.06^{+0.02}_{-0.02}$	$4.4^{+0.1}_{-0.1}$	19^{+2}_{-2}	459	χ^2 too high to trust fit
004425.73+412242.4	$4.6^{+0.6}_{-0.7}$	$3.4^{+0.5}_{-0.3}$	$4.4^{+0.2}_{-0.2}$	$3.8^{+0.2}_{-0.2}$	14^{+11}_{-8}	8	B
004448.13+412247.9	$5.7^{+0.1}_{-0.1}$	$3.9^{+0.1}_{-0.1}$	$4.63^{+0.02}_{-0.02}$	$4.6^{+0.1}_{-0.1}$	49^{+4}_{-3}	940	high χ^2, flat SED
004518.38+413936.6	$2.8^{+0.5}_{-0.4}$	$4.1^{+0.2}_{-0.2}$	$4.2^{+0.1}_{-0.1}$	$1.4^{+0.6}_{-0.6}$	4^{+2}_{-1}	10	B
004527.88+413905.5	$6.5^{+0.1}_{-0.1}$	$3.8^{+0.1}_{-0.1}$	$4.70^{+0.03}_{-0.04}$	$5.2^{+0.1}_{-0.1}$	158^{+39}_{-25}	667	χ^2 too high to trust fit
004528.24+412943.	$5.4^{+0.1}_{-0.1}$	$2.5^{+0.1}_{-0.1}$	$4.28^{+0.02}_{-0.02}$	$2.8^{+0.1}_{-0.1}$	26^{+2}_{-2}	3409	χ^2 too high to trust fit

Note. Output best-fit parameters for stellar optical counterparts using the BEAST SED fitting code for all point sources listed in Table 2. Median values ±33% are listed. After the posterior distribution is complete, the BEAST code calculates the χ^2 value for the most likely model, listed here. Fits are considered robust for $\chi^2 < 12$. The *Probable Spectral Type* column lists the most likely spectral type of each source, based on its best-fit physical properties. See Section 4.2 for a more detailed description of BEAST SED fitting.

green, and blue correspond to the soft, intermediate, and hard accretion states, respectively).

In these diagrams we also include accreting B*e*-XRB pulsars with available RXTE-PCA spectra (e.g., Reig 2011) following the same procedure as for the black hole X-ray binaries. Their intrinsically hard X-ray spectra clearly separate them even from the locus of the hard-state black hole X-ray binaries (A. Zezas et al. 2018, in preparation). Finally, we include spectra of Z-track neutron star Low-mass X-ray binaries (LMXBs).

Determining background AGN contamination is difficult as their hard X-ray colors and luminosities can be similar to those of compact objects in the disk of M31 (see e.g., Tozzi et al. 2006). While most background galaxies do not have *NuSTAR* classifications because they do not have enough counts in the hard (12–25 keV) band, they can occupy similar regions of the hardness ratio and hardness-intensity diagrams as compact objects in M31. Two sources that were determined to be background galaxies using PHAT imaging are plotted as black circles outlined in orange in Figure 5. This highlights the importance of incorporating data at optical wavelengths to remove these sources from our hard X-ray sample.

4.2. SED Fitting of Stellar Optical Counterparts

We obtained SED fits of 15 point-source optical counterparts using the code BEAST (Gordon et al. 2016). The BEAST code fits the observed SED of an individual star in M31 with theoretical SEDs from the Padova stellar evolution models (Marigo et al. 2008) using a Bayesian statistical approach. The code assumes single-star evolution and that sources are in M31. Photometric bias and uncertainty are applied from artificial star tests performed on the data. The input for the BEAST code is the six-band photometry and artificial star tests of the star measured by the PHAT survey. The code uses upper limits as constraints. Output parameters include several primary and derived quantities. Primary fit outputs include initial stellar mass, $A(V)$ (dust extinction), and stellar metallicity. Derived quantities include luminosity, effective temperature, and stellar surface gravity. Output physical parameters for the 15 point-source optical counterpart candidates in our sample are listed in Table 3.

As part of the fitting, χ^2 values are computed assuming multi-variate Gaussian uncertainties, either uncorrelated or correlated. The probability of a given model is proportional to χ^2, letting us use the χ^2 value as a relative assessment of the "goodness of fit."

In most HMXBs, the donor star is much brighter than the accretion disk at optical wavelengths, so the fits with low χ^2 values should be robust. The BEAST code is designed to fit individual stars, and so it will return a poor fit if a point source is not an individual star. Examples of systems that might return a poor fit include parts of a multiple star system, background AGN, companions in XRBs that have been irradiated by their associated X-ray source (Phillips & Podsiadlowski 2002), stars contaminated with light from the compact object's accretion disk, chance superpositions of sources, or B*e* star donors with a red excess from the accretion disk relative to the underlying B-star spectrum.

Table 3 shows a clear division in χ^2 values: $\chi^2 \lesssim 12$ or $\chi^2 \gtrsim 50$. Examples of these two categories are shown in Figure 6. The lower χ^2 fit appears similar to a stellar SED model, while the high χ^2 appears to have a flat SED. Based on the clear division in fit quality as well as SED appearance, we decided that fits with $\chi^2 \lesssim 12$ likely have SEDs consistent with stars in M31, while fits with higher χ^2 values do not. Thus, we did not determine a spectral type for sources with χ^2 values above ∼12.

Table 3 lists the probable spectral type given the best-fit physical parameters. We determine masses, temperatures, and luminosities for 7 of the 15 point sources that are consistent with a B-type star, and therefore very strong HMXB candidates. B-type stellar classification was determined for $4\,M_\odot \lesssim M \lesssim 17\,M_\odot$ and $4.0\,\mathrm{K} \lesssim \log(T_{\rm eff}) \lesssim 4.5\,\mathrm{K}$ (e.g., Silaj et al. 2010). Figure 6 shows the SED fit for the optical

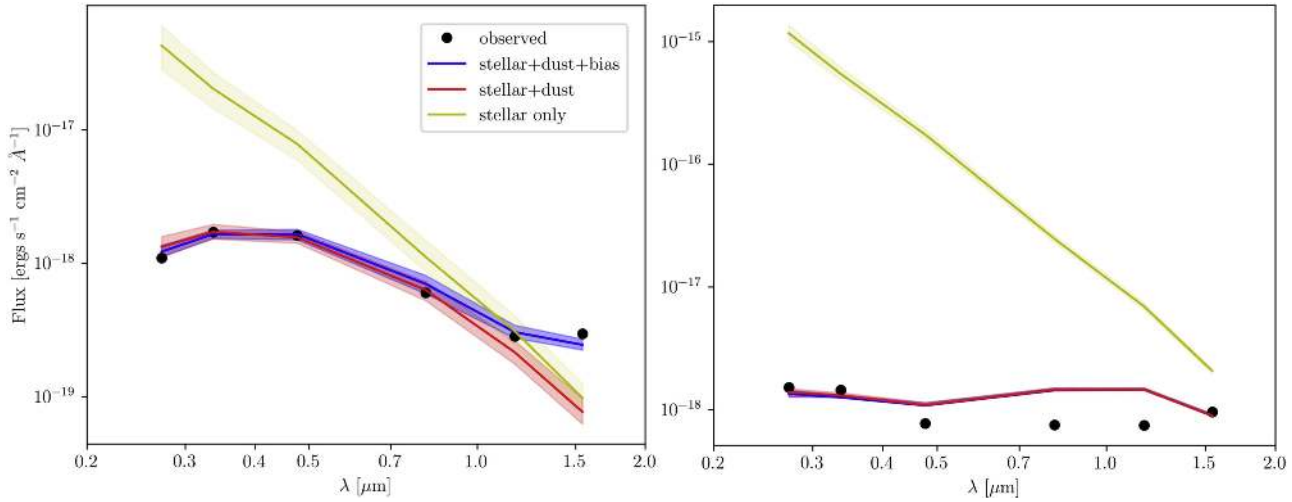

Figure 6. BEAST SED fits for the optical counterparts of two HMXB candidates, 004321.48+411556.9 (left) and 004448.13+412247.9 (right). The counterpart in the left panel has a robust fit and is most likely a B-type star. The counterpart in the right panel returns a poor fit. Black points show measured photometry for the optical counterparts from the PHAT data set, listed in Table 2. The colored lines show the median fit ±33% errors of the three different models. Yellow shows the stellar-only model, red shows a stellar+dust model, and blue shows a stellar+dust+bias model. The observational bias is determined using artificial star tests.

counterpart to 004321.48+411556.9, an example of a good fit for a B-type star.

One optical counterpart (004249.22+411815.8) is classified as a possible red helium-burning star given its high luminosity, low temperature, and low surface gravity. The hard X-ray source associated with this optical counterpart does not have a *NuSTAR* classification.

We do not rule sources out as HMXB candidates due to fits because a poor fit may be returned for stars that have been irradiated by their associated X-ray source or contaminated with light from their compact object's accretion disk, as discussed previously in this section. However, sources with good fits to stars in M31 may be stronger than those that do not.

We find two point sources that have relatively flat SEDs, noted in Table 3. Figure 6 illustrates the BEAST fit for the optical counterpart of one of these sources, 004448.13 +412247.9. The BEAST attempts to fit the point source as a hot star with a high $A(V)$ (e.g., 4.6 mag of extinction, as shown in Figure 6, compared to the 1.2 mag of dust extinction shown for a star that has a robust fit to a B-type star) and still returns a poor fit. The other source that has an optical counterpart with a flat SED is 004316.11+411841.5. Such a flat SED may be indicative of a background AGN.

4.3. SFHs of HMXB Candidates

We used the spatially resolved recent SFH of M31 by Lewis et al. (2015) to determine likely ages of HMXB candidates in our sample. Lewis et al. (2015) inferred these SFHs using CMDs of 100 pc by 100 pc regions in the M31 disk.

We assume that HMXBs contain secondary stars more massive than 7 $M_{\odot}$, which have lifetimes of 10 Myr. Thus we conservatively restrict our age distribution analysis to < 60 Myr. The time resolution of the SFHs is log(time) = 0.1 year. SFHs are not available for regions too close to the bulge of M31 (in PHAT bricks 1 and 3) because crowding does not allow for accurate CMD fitting, and so not all HMXB candidates are included in our analysis. For that reason, 8 of the 15 HMXB candidates are used in the SFH analysis: 004335.91+411433.4, 004350.76 +412118.1, 004404.75+412127.2, 004404.75+ 412127.2, 004425.73+412242.4, 004448.13+412247.9, 004518.38+ 413936.6, 004527.88+413905.5, and 004528.24+412943.9.

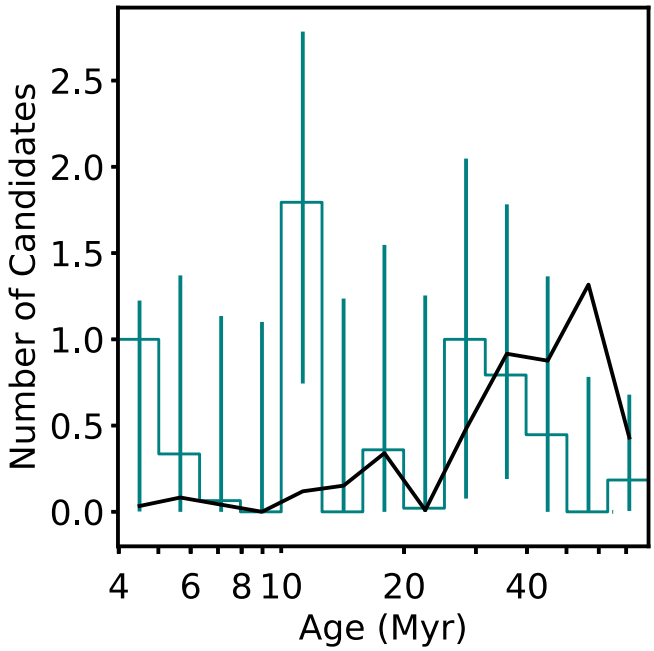

Figure 7. Histogram showing the number of HMXB candidates associated with each time bin based on their spatially resolved star formation histories from Lewis et al. (2015). The time bins are defined by the age resolution of the SFHs. The black line represents the average of 100 random samples of regions in M31 not associated with HMXB candidates, providing a reference for the overall SFH of M31.

For each HMXB region containing an HMXB candidate, we calculated the total stellar mass formed in the past 60 Myr. For each time bin younger than 60 Myr, we calculated the fraction of the mass formed in that bin. This fraction gives the normalized probability that the given HMXB candidate formed in that time bin. We take the uncertainties in the SFH into account by sampling the SFH 1000 times and recalculating the age distribution. We take the 16th and 84th percentile in each time bin to determine uncertainties. The number of HMXB

candidates expected to form in each time bin is shown in Figure 7 in teal.

We then compare the probability distribution for regions with HMXB candidates to the rest of the M31 disk. We do this by randomly selecting 8 regions from a sample of ∼49 regions containing known background galaxies identified by Williams et al. (2018, submitted) in the PHAT bricks observed by *Chandra* and *NuSTAR* and immediately adjacent. We use the SFH for regions around background galaxies because these should be randomly distributed throughout the disk and not correlated with the HMXB population. We perform this random selection 100 times and plot the average expected number of candidates in each time bin in black in Figure 7.

Using the subsample of eight HMXB candidates with SFHs, we were able to determine that about three HMXBs in our sample have an age of ∼25–50 Myr, 2 are ∼10 Myr old, and 1 is ∼4 Myr old. Two of the HMXB candidates analyzed, 004350.76+412118.1 and 004404.75+412127.2, are found in regions without significant star formation in the last 60 Myr, making them weaker HMXB candidates.

The ages of candidates between 25 and 50 Myr are fairly consistent with random draws from the disk of M31. This indicates that the regions with these HMXB candidates do not appear to be a different age than the average population. The peak in star formation in regions surrounding HMXB candidates in the 10–12 Myr and 4 Myr time bins are more significant deviations from the overall SFH of M31, as demonstrated in Figure 7. These may be probing the prompt HMXB formation channel in M31.

5. Discussion

Our measurements allow many detailed comparisons of the X-ray sources in this region of M31. First, we can compare the optical and X-ray characteristics of the sources with counterparts. Next, we can compare those characteristics with the age distribution of the surrounding stellar populations as an additional consistency check, and finally we can consider the sources in globular clusters to look for X-ray characteristics unique to that specific subclass. We discuss all of these comparisons below.

5.1. *Comparing* NuSTAR *and* HST *Source Classification*

In Figure 8 we compare the classification of the compact object determined by *NuSTAR* colors with the type of the associated optical counterpart. Sources that fall within the "none" *NuSTAR* classification did not have enough counts in all three X-ray bands to be accurately classified. Sources with the "none" optical counterpart classification did not have a clear optical counterpart in PHAT imaging.

Nine *NuSTAR* classified non-magnetized neutron stars in our sample are found within globular clusters, and we find no pulsars or hard-state black holes in globular clusters. Roughly equal numbers of pulsars, non-magnetized neutron stars, and hard-state black holes have point sources as optical counterparts. We also find that four pulsars in our sample are HMXB candidates with point-source optical counterparts, while four have no optical counterpart. Pulsars without optical counterparts could be part of a low- or intermediate-mass X-ray binary system. This suggests that the pulsars in our sample are not preferentially in HMXB systems. When we compare the *NuSTAR* source classifications with the results of the BEAST SED fitting (summarized in Table 4), we note that none of the HMXB candidates with classified hard-state black holes have good SED fits to B-type stars.

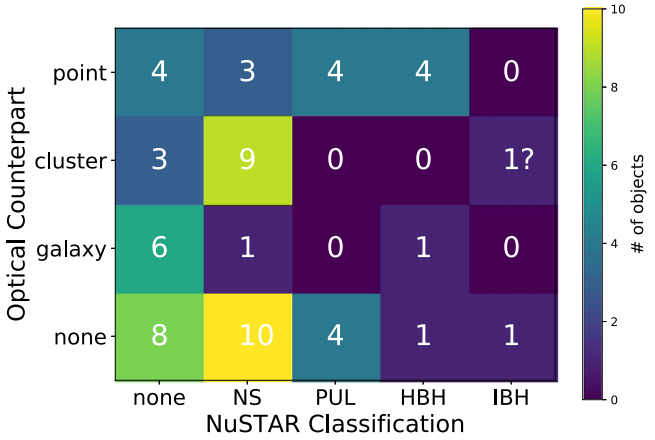

Figure 8. Comparing the *NuSTAR* X-ray classification and optical counterpart of sources observed with *NuSTAR*, *Chandra*, and PHAT. *NuSTAR* classifications are defined as follows: NS = non-magnetized neutron star, PUL = pulsar, HBH = hard-state black hole, and IBH = intermediate-state black hole. For a further discussion of the *NuSTAR* classified intermediate-state black hole associated with a cluster, see Section 5.5.

Table 4
HMXB Candidate Classification Data

Catalog Name	Good BEAST Fit	SFH Indicates Young SF	*NuSTAR* Classified HBH/Pulsar/NS
004240.31+411845.6		···	✓ (hbh)
004249.22+411815.8	✓	···	
004308.63+411248.4	✓	···	✓ (pul)
004316.11+411841.5		···	✓ (hbh)
004321.08+411750.6	✓	···	✓ (ns)
004321.48+411556.9	✓	···	
004335.91+411433.4	✓	✓	
004339.06+412117.6	✓	···	✓ (pul)
004350.76+412118.1			✓ (ns)
004404.75+412127.2			✓ (hbh)
004425.73+412242.4	✓	✓	✓ (pul)
004448.13+412247.9		✓	✓ (ns)
004518.38+413936.6	✓	✓	
004527.88+413905.5		✓	✓ (pul)
004528.24+412943.9		✓	✓ (hbh)

Note. Table evaluating likelihood of HMXB candidates based on BEAST fits, SFH, and *NuSTAR* classification of compact objects. Check marks are given when a source is likely to be an HMXB using the given criteria: the source has a good BEAST SED fit to a stellar companion, the source is in a region with young (within the last 60 Myr) star formation, or the *NuSTAR* classification of the compact object is a pulsar, non-magenetized neutron star, or a pulsar. Ellipses in the SFH column indicate that no star formation history is available for the region around the source because of crowding near the bulge of M31.

It is important to note *NuSTAR* sources 57 and 70, which are each blends of two *Chandra* sources (see Section 3.1). In both cases, one *Chandra* source has a globular cluster optical counterpart and the other has no optical counterpart. Source 70 is classified as a non-magnetized neutron star, suggesting that the *Chandra* source associated with a globular cluster may dominate the light detected by *NuSTAR*. Source 57 is classified as an intermediate-state black hole, for which it would be unusual if it were associated with a globular cluster. We

compared the 4–8 keV *Chandra* flux of the two *Chandra* sources that matched *NuSTAR* source 70 and 57. The energy fluxes were consistent within errors, so we could not determine which source dominated the flux observed by *NuSTAR*. The *NuSTAR* source classification may be affected by the blend of the two *Chandra* sources, and thus additional investigation is needed to confirm if the source classification is a result of the blend.

We compare PHAT imaging with *NuSTAR* classifications to remove background AGN contamination from our sample. We find two cases where sources classified as compact objects in the disk of M31 were determined to be background galaxies using PHAT imaging. See Figure 9 for PHAT images used to identify these background galaxies. The apparent misclassification of these two sources does not affect the conclusions of this paper, but illustrates the power and necessity of combining *NuSTAR* observations with *HST* data to eliminate background AGN contamination. These two sources are 004527.30+413254.1 (*NuSTAR* source 118, discussed in more detail in Section 5.5), which is classified as a hard-state black hole, and 004530.61+413600.4 (*NuSTAR* source 116), which is classified as a non-magnetized neutron star. Both sources have resolved background galaxies as optical counterparts in the PHAT imaging.

5.2. Evaluation of HMXB Candidates

Table 4 summarizes our investigation of HMXB candidates, identified by selecting for hard X-ray sources spatially coincident with UV-bright point-source optical counterparts. We evaluate whether a source is a likely HMXB using three methods: (1) SED fitting with the BEAST code to determine if a massive, young star appears to be the donor, (2) age estimation using spatially resolved SFHs, and (3) compact object classification using *NuSTAR* hard X-ray colors and luminosities.

We consider any hard X-ray source with a UV-bright point-source optical counterpart an HMXB candidate, even if it does not satisfy all three criteria. For example, the optical counterpart could have a poor SED fit because of irradiation from the compact object or mass transfer, as discussed in Section 4.2. Additionally, not having a *NuSTAR* compact object classification does not rule out an HMXB candidate. Sources must have enough flux in all three *NuSTAR* bands to be classified, so a source could remain unclassified if it is too faint or there is too much absorption to be detected in all bands. We comment that having a *NuSTAR* compact object classification or good optical companion SED fit makes an HMXB a stronger candidate because we have more information about the system.

We find that HMXB candidate 004425.73+412242.4 satisfies all three criteria. It has a good SED fit indicating a B-type donor star and the SFH in the region around this source shows star formation bursts within the last 60 Myr. The *NuSTAR* colors and luminosities of this source indicate that it is likely a pulsar.

Three HMXB candidates in our sample, 004308.63+411248.4, 004321.08+411750.6, and 004339.06+412117.6, have optical counterparts that have good SED fits to B-type stars and have a *NuSTAR* classification of pulsar or non-magnetized neutron star. However, because they are located close to the bulge of M31, spatially resolved SFHs are not available for the region surrounding these candidates.

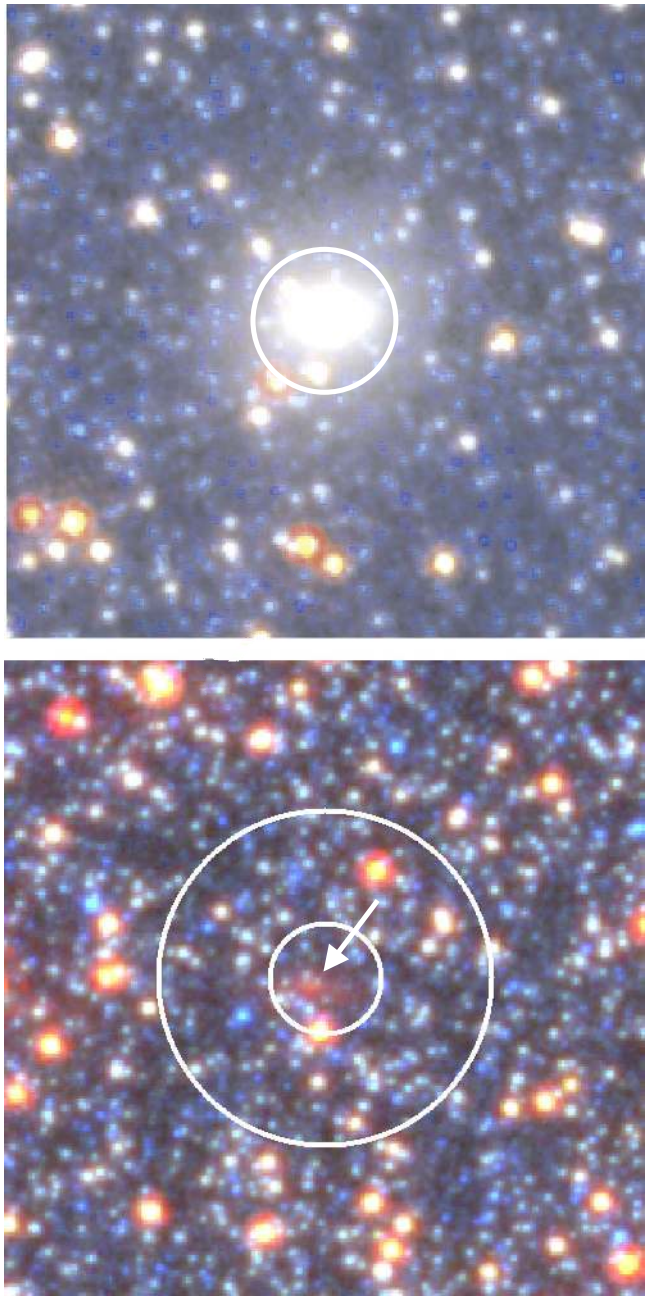

Figure 9. Images showing background galaxies associated with *NuSTAR* source 118 (left) and 116 (right). These color images were created with *HST* imaging from the PHAT survey and use the F160W filter as red, F814W as green, and F336W as blue. Images are 10″ on a side. White circles indicate the 1 and 3σ *Chandra* positional errors. The 1σ error is not visible on the image on the left as it lies on top of the bright galaxy. The galaxy on the left is bright, extended, and elliptical in shape. The galaxy on the right is much fainter. It is visible as a faint, red, extended source within the 1σ error error circle, indicated with an arrow. For more discussion of background galaxies, see Section 5.1.

Four HMXB candidates in our sample, 004448.13+412247.9, 004518.38+413936.6, 004527.88+413905.5, and 004528.24+412943.9, satisfy two of the three criteria listed in Table 4. All of these sources are found in regions with recent star formation, but are either lacking a good SED fit to a B-type star or do not have a compact object classification from *NuSTAR*.

Four HMXB candidates in our sample have either a good SED fit for a B-type donor star (004249.22+411815.8 and

004321.48+411556.9) or *NuSTAR* classified compact object (004240.31+411845.6 and 004316.11+411841.5). No SFHs are available for these sources because of their proximity to the bulge of M31.

Two HMXB candidates in our sample lack good SED fits to B-type companion stars and are found in regions with no significant star formation in the last 60 Myr. These sources are indicated as HMXB candidates because they have UV-bright point sources associated with a *NuSTAR*-detected hard X-ray source. Source 004350.76+412118.1 is classified as a non-magnetized neutron star and source 004404.75+412127.2 is classified as a hard-state black hole. The lack of star formation and poor SED fits could indicate that these sources are background AGN. Thus, further multiwavelength observations are needed.

5.3. Comparison with Other Hard X-Ray Observations of M31

Three sources in our sample were investigated in detail by Stiele & Kong (2018) in a *NuSTAR* survey of the central region of M31. This study was designed to overlap with a previous *XMM-Newton* survey of M31 (Stiele et al. 2011). Stiele & Kong (2018) use hardness ratios to classify four X-ray sources previously classified only as "hard" using *XMM-Newton* data as X-ray binaries. Two of these new XRB candidates are in our sample: *NuSTAR* sources 47 and 65. We classify these sources as non-magnetized neutron stars and find no optical counterpart, suggesting that they could be LMXBs because these types of stars would be too faint to be observed in the PHAT survey. Additionally, Stiele & Kong (2018) discuss several sources that are too hard to be located in the XRB area of their hardness-ratio diagrams. One of these sources is in our sample (*NuSTAR* source 91) and is classified as a pulsar with no optical counterpart.

5.4. Ages of Stellar Populations Hosting HMXBs

The ages (20–50 Myr) of the regions surrounding most HMXB candidates in this study are consistent with the results seen in the Small Magellanic Cloud (SMC; Antoniou et al. 2010). In the SMC, Be/X-ray binaries are the most numerous subclass of HMXBs, and they are found in regions with star bursts that occurred between 25 and 60 Myr ago. In the LMC, X-ray binaries were found associated with younger regions, between 6 and 25 Myr old (Antoniou & Zezas 2016). The statistically significant increase in the number of HMXB candidates we found in M31 in regions with a star formation burst 10 Myr ago (Figure 7) aligns with the young ages found in the LMC.

We examined the three sources in regions with a strong peak in star formation rate in the 10 Myr bin, since this is a 1.5σ deviation from the background population (as shown in Figure 7). The stellar population surrounding one source in particular, 004425.73+4122241.8, experienced almost all of its star formation in the 10 Myr time bin. This HMXB candidate has been classified as a pulsar with a B-type stellar companion, determined by its *NuSTAR* colors and BEAST SED fit. Two other sources are also located in regions with SFR peaks in this time bin: 004518.38+413936.6 and 004448.13+412247.9. Source 004518.38+413936.6 has an optical counterpart that is classified as a B-type star with no *NuSTAR* classification for the compact object. Source 004448.13+412247.9 is classified as a neutron star by *NuSTAR*, but the BEAST SED fit quality is too low to determine the spectral type of the companion, but it is probably not a single star in M31.

Source 004518.38+413936.6 (with a B-type stellar optical counterpart) is located in a region that also experienced significant star formation in the 4 Myr time bin. Note that connecting HMXB populations to stellar ages is important to constrain formation models of compact objects. Rappaport et al. (2005) and Justham & Schawinski (2012) predict that a time delay of 10 Myr (assuming instantaneous burst of star formation) or 200 Myr (continuous star formation) may be expected between the onset of star formation and the production of X-rays, depending on SFH. XRB pulsars have been found with similar ages in the Magellanic Clouds. Li et al. (2016) found an X-ray pulsar with an O-type counterpart star in the SMC, suggesting that the system is ∼5–6 Myr old, and Belczynski et al. (2008) find that XRB pulsars can form at ages as young as ∼5 Myr. HMXBs associated with very young stellar ages (10 Myr or younger) but with B-star secondaries can place a particularly important constraint on initial mass ratios of HMXBs, as such an object must have had a much more massive companion with a lifetime short enough to have become the accreting compact object.

5.5. X-Ray Sources in Clusters

Maccarone et al. (2016) investigated hard X-ray sources in globular clusters using combined *Swift-NuSTAR* spectroscopy. Our sample of 64 X-ray sources observed by *NuSTAR*, *Chandra*, and *HST* includes four of the five sources in that study. These sources are not HMXB candidates as they were found to be spatially coincident with globular clusters, not point sources. We find that three of these sources (*NuSTAR* sources 65, 104, and 120 in our sample) are classified as neutron stars and one (*NuSTAR* source 118) is classified as a hard-state black hole.

We determine the hard-state black hole has a background galaxy optical counterpart rather than a globular cluster, using its PHAT imaging. We also cross-reference the globular cluster and background galaxy catalogs published by the PHAT survey and find that this source is classified as a galaxy based on its morphology (Johnson et al. 2015).

This source was investigated in detail by Dorn-Wallenstein et al. (2017), who found that it has a spectroscopic redshift, which agrees with our classification as a background galaxy. This source highlights the importance of incorporating data at optical wavelengths to remove background AGN contamination.

We identify one source (*NuSTAR* source 57) classified as an intermediate-state black hole that may be associated with a globular cluster. Some caution is warranted in interpreting this source because it is associated with two separate *Chandra* sources, 004255.61+411834.8 and 004255.19+411835.7, and hence the *NuSTAR* spectrum is probably a superposition of two different source spectra. The former of the two *Chandra* sources has a globular cluster optical counterpart, while we do not see evidence for a globular cluster associated with the latter source.

The potential connection between the globular cluster and a black hole is intriguing. For quite some time, it was thought that the Spitzer (1969) instability would lead to mass segregation that would, in turn, expel most or all stellar mass black holes from globular clusters from globular clusters (Kulkarni et al. 1993; Sigurdsson & Hernquist 1993). The discoveries of strong candidate globular cluster black holes in

external galaxies (e.g., Maccarone et al. 2007) and in the Milky Way (e.g., Strader et al. 2012; Chomiuk et al. 2013; Giesers et al. 2018) has helped motivate and support theoretical work that has shown that globular cluster may retain black holes (e.g., Mackey et al. 2008; Sippel & Hurley 2013; Morscher et al. 2015).

The globular cluster G1 is of special interest as it has been suggested to contain an intermediate-mass black hole (IMBH) on the basis of stellar dynamical evidence (Gebhardt et al. 2002; Baumgardt et al. 2003, for an alternative view). Its X-ray source is consistent with accretion from the putative IMBH (Pooley & Rappaport 2006). It appeared as a detectable radio source in VLA data (Ulvestad et al. 2007), but sensitive radio data later found only deep upper limits (Miller-Jones et al. 2012), again providing an ambiguous determination of whether the cluster contains an IMBH. The IMBH classification is highly uncertain, as the X-ray observations are also consistent with emission from an LMXB (Kong et al. 2010; Miller-Jones et al. 2012). Potentially, deep *NuSTAR* imaging could provide some additional clues about this interesting globular cluster source as well.

Still, the total number of strong candidate black holes in globular clusters remains relatively small, especially at distances where the structural parameters of the clusters are measurable, which means that *NuSTAR* source 57 in our sample merits follow-up work to further test the black hole hypothesis.

6. Conclusions

In this work we present 15 HMXB candidates: hard X-ray sources observed by *NuSTAR* and *Chandra* that are spatially associated with UV-bright point sources from the PHAT catalog.

We investigated the correlation between the *NuSTAR* determined compact object type and the optical counterpart determined with PHAT imaging. We find nine *NuSTAR* classified non-magnetized neutron stars associated with star clusters, making this the strongest correlation in our sample, and this also agrees with the findings in Maccarone et al. (2016).

We did not find any pulsars or hard-state black holes associated with star clusters. There did not appear to be a preference for non-magnetized neutron stars, pulsars, or hard-state black holes associated with UV-bright point-source optical counterparts. None of the HMXB candidates in our sample with hard-state black hole compact objects have a companion star with a good SED fit to a B-type star.

We also find an equal number of pulsars in HMXB and LMXB systems. For the pulsars, this may point toward an interesting result, but our source statistics are too small to tell; further observations are needed. However, either the pulsars are not HMXBs and might have intermediate donor masses such as those found in other M31 pulsar systems (e.g., Esposito et al. 2016; Yukita et al. 2017), or perhaps their pulsar identifications are not as secure.

We determined likely ages for HMXB candidates using published SFHs. We find that three HMXBs in our sample are associated with stellar populations between 25 and 50 Myr old, and two to three HMXB candidates are associated with younger stellar populations: one to two are ∼10 Myr old, and one is ∼4 Myr old. These ages agree with findings in the Magellanic Clouds, M33, NGC 300, and NGC 2403. The ages we find in M31 and those found in other galaxies suggest two potential formation channels for HMXBs.

Beyond our results from investigating individual X-ray sources, this study demonstrates the ability to study both the compact object and companion star in an XRB from the hard X-rays to the near-IR using *NuSTAR*, *Chandra*, and *HST*. We were able to use classifications by Wik et al. (2018, in preparation) of hard X-ray sources as neutron stars or black holes based on their X-ray colors and luminosities. Matching the *NuSTAR* sources to *Chandra* allowed us to determine the positions of these X-ray sources with increased accuracy, and thus find and classify their optical counterparts using the PHAT data set. This study is an exciting foray into the combination of hard X-ray and deep optical observations in nearby galaxies. Given the maturity of the PHAT data set, we are able to harness the data products created by the many scientists on the PHAT team to determine ages and spectral types.

We look forward to continuing this work in other local galaxies as more deep *HST* and *NuSTAR* observations are made. We also plan to compare these observational XRB population data to the predictions of theoretical population synthesis codes (e.g., Sørensen et al. 2017) to place constraints on models of the formation and evolution of these systems.

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